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# Dark Energy with Spectroscopic Observations of the Universe

Julian E Bautista

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# Habilitation à Diriger des Recherches

Présentée à Aix-Marseille Université  
le 24 octobre 2023

## Dark Energy with Spectroscopic Observations of the Universe

Julián Ernesto BAUTISTA

### Discipline

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### Spécialité

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# Abstract

This manuscript summarises my research over the past 12 years in the field of observational cosmology, as well as the work by early career scientists for whose I was the main or the co-supervisor. The main topic of our work is the study of dark energy with spectroscopic galaxy data from the Sloan Digital Sky Survey and the Dark Energy Spectroscopic Instrument. We explored three quite different regimes in terms of redshift ranges and types of data: Ly $\alpha$  forests at high-redshift ( $2 < z < 3.5$ ), galaxies at mid-redshift ( $0.6 < z < 1.0$ ) and type-Ia supernovae at low-redshift ( $0 < z < 0.1$ ). We mostly focused on the statistical properties of these samples, estimating two-point functions and measuring the scale of baryon acoustic oscillations and the effect of redshift-space distortions. We explored new techniques for measuring weak-lensing from Ly $\alpha$  forests and the link between galaxies and the 21cm radio signal. All this research is linked to the challenge of precisely measuring the expansion rate of the Universe and the growth-rate of structures, with the hope of finding deviations from the standard  $\Lambda$ CDM cosmological model based on general relativity.

Keywords: cosmology, dark energy, spectroscopy





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# Contents

<b>Abstract</b>	<b>iii</b>
<b>Acknowledgements</b>	<b>v</b>
<b>1 Introduction</b>	<b>1</b>
1.1 The history of our Universe and its big open questions . . . . .	1
1.2 Evidences for the accelerated expansion of the Universe . . . . .	2
1.3 Model of an expanding Universe . . . . .	3
1.3.1 Assumptions and ingredients . . . . .	3
1.3.2 The expansion rate . . . . .	3
1.3.3 Distances in cosmology . . . . .	4
1.3.4 Dark energy models . . . . .	6
1.4 Model of the large-scale structures . . . . .	6
1.4.1 Statistical description of perturbations . . . . .	7
1.4.2 Configuration and Fourier space . . . . .	7
1.4.3 Cosmological dependency of the power spectrum . . . . .	8
1.4.4 The amplitude of the power spectrum . . . . .	10
1.5 Cosmological probes of expansion . . . . .	11
1.5.1 Direct measurements of $H_0$ . . . . .	11
1.5.2 Type-Ia supernovae . . . . .	12
1.5.3 Big Bang nucleosynthesis . . . . .	13
1.5.4 Baryon acoustic oscillations . . . . .	13
1.5.5 The cosmic microwave background . . . . .	15
1.5.6 Redshift-space distortions . . . . .	16
1.5.7 Weak gravitational lensing . . . . .	17
References . . . . .	19
<b>2 Observing the Universe with spectroscopy</b>	<b>25</b>
2.1 Selecting the objects to observe . . . . .	25
2.2 Pointing fibres to the sky . . . . .	26
2.3 From electrons to spectra . . . . .	28
2.4 From spectra to redshifts . . . . .	31
References . . . . .	32
<b>3 The high-redshift Universe and its forests</b>	<b>35</b>
3.1 Forests as a tracer of the Universe's structures . . . . .	35
3.2 Baryon acoustic oscillations in the forests . . . . .	38
3.2.1 Transmission field . . . . .	38
3.2.2 Two-point correlation functions . . . . .	40

3.2.3	Correction matrices . . . . .	40
3.2.4	The model . . . . .	41
3.2.5	BAO constraints . . . . .	43
3.3	Impact of redshift errors . . . . .	44
3.4	Weak-lensing of forests . . . . .	47
	References . . . . .	50
<b>4</b>	<b>The mid-redshift Universe and its galaxies</b>	<b>55</b>
4.1	From redshifts to clustering catalogues . . . . .	55
4.2	Reconstruction of linear density field . . . . .	59
4.3	Mock catalogues . . . . .	61
4.4	From catalogues to clustering estimates . . . . .	61
4.4.1	Correlation function . . . . .	62
4.4.2	Power spectrum . . . . .	62
4.4.3	The clustering of eBOSS DR16 LRGs . . . . .	64
4.5	Baryon acoustic oscillations . . . . .	65
4.6	Redshift-space distortions . . . . .	67
4.7	Joint clustering analysis in Fourier and Configuration space . . . . .	71
4.8	Cross-correlation with radio surveys . . . . .	73
	References . . . . .	76
<b>5</b>	<b>The low-redshift Universe and its velocities</b>	<b>83</b>
5.1	Measuring peculiar velocities . . . . .	83
5.1.1	The Tully-Fisher relation of spiral galaxies . . . . .	84
5.1.2	The Fundamental Plane of elliptical galaxies . . . . .	84
5.1.3	Type-Ia supernovae . . . . .	85
5.2	Biases affecting velocities . . . . .	86
5.3	Growth-rate measurements with velocities . . . . .	88
5.3.1	Maximum likelihood method . . . . .	88
5.3.2	Compressed two-point statistics . . . . .	90
5.3.3	Density-velocity comparison . . . . .	91
5.3.4	Forward-model of density and velocity field . . . . .	92
5.4	Current measurements . . . . .	92
5.5	DESI and ZTF . . . . .	92
5.6	Growth-rate measurement from simulated ZTF data . . . . .	94
5.6.1	Simulating ZTF type-Ia supernovae with peculiar velocities . . . . .	95
5.6.2	Measuring peculiar velocities and the growth-rate . . . . .	97
	References . . . . .	99
	<b>Conclusion and future outlook</b>	<b>105</b>

# Chapter 1

## Introduction

### 1.1 The history of our Universe and its big open questions

Gathering all the knowledge in physics humanity could learn to this day, we are able to come up with a quite interesting and scientifically accurate story, describing how our Universe evolved from more than 13.8 billions years ago to nowadays. Our story can explain most of what we observe in the sky, while being consistent with essentially all experimental results obtained on Earth and its Solar neighbourhood. However, the story of our Universe has a few but quite important weak spots. Either because where we do not have a satisfactory physical explanation for some observed phenomena, or simply because some predicted phenomena are not currently observable (and potentially will never be). That also explain my use of the word *story* instead of *history*, which would only contains verified facts.

The story in a nutshell goes as follows: the “beginning” of the Universe<sup>1</sup>, also known as the *Big Bang* is the start of the *inflation*, an epoch of exponential expansion of space that we believe happened around 13.8 billion years ago. Inflation transformed quantum, microscopic fluctuations of space into macroscopic density fluctuations of the Universe’s constituents: quarks, photons, neutrinos and dark matter. After inflation, the Universe continues to expand but much slowly. The average temperature of the Universe decreases adiabatically. Quarks start to gather to form protons, neutrons and relics of heavier atoms, such as deuterium, tritium, helium, lithium and so on. All this during the first few minutes of the Universe’s history, in the so called *Big Bang nucleosynthesis*. After 380 thousand years, electrons rebind with protons to form neutral hydrogen, while photons scatter for their last time with baryons as the Universe becomes more diffuse. The photons from this epoch are observable today and are known as the *cosmic microwave background*. The Universe then enters its *dark ages*, where hydrogen is mostly neutral and stars have not yet formed. After several million years, gravity clumps hydrogen into dense clouds that reach high enough temperatures to start thermonuclear reactions in their core. The first stars are formed. Then the first galaxies. These galaxies merge into larger galaxies, galaxies form small virialised groups, large groups, clusters, super clusters. Clusters, filamentary structures and voids compose the so called cosmic web. At around half of the Universe’s age, the expansion starts to speed up again, accelerating, as if gravity became repulsive on very large scales. Then, on one of the many billions of galaxies,

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<sup>1</sup>The actual beginning of the Universe can be seen as a complex philosophical question or simply as an ill defined physical concept. This discussion is beyond the scope of this work and/or my capabilities. Here, I simply name as “beginning” the first epoch of the Universe for which we have a widely known physical theory to describe it.

the *Milky Way*, planet Earth formed around a somewhat isolated star, the Sun, and here we are today, observing the sky, trying to explain the whole Universe.

This story is based on a few assumptions: that general relativity (GR) describes gravity; that the Universe can be modelled as statistically homogeneous and isotropic on large enough scales; and that the Universe is composed of the standard model particles and interactions, with the addition of two exotic components: dark energy and dark matter.

The origin or nature of the dark sector is one of the biggest mysteries in physics. Several models of dark matter are based on particles beyond the standard model which have not yet been detected, despite the huge experimental efforts towards this goal. The simplest model of dark energy is a cosmological constant which is actually part of GR equations. The possibility that dark energy could be an exotic fluid or the manifestation of an extra force has predicted implications that were not yet observed either.

Another weak spot of the story of our Universe is inflation. Up to now, inflation is a successful theory explaining an important set of observations: the homogeneity of the photon temperature from the cosmic microwave background on scales larger than the horizon; the flatness of space; and the seeds of structure formation. However, the predicted and observable effects of inflation were not yet observed, such as primordial gravitational waves.

Solving for the mysteries of the dark sector and inflation would represent major breakthroughs in physics. Large teams of scientists, including myself, are dedicated to these puzzles, either theoretically or experimentally.

My past work has been focused on observations related to the understanding of dark energy. The rest of this chapter will be dedicated to explaining the problem of accelerated expansion, the dark energy model, alternative solutions, and how to observe the expansion.

## 1.2 Evidences for the accelerated expansion of the Universe

Not long after developing his theory of gravity, Einstein wanted to write a model for a static Universe filled with standard matter. Since gravity is attractive, his solution was to include a new constant term,  $\Lambda$ , in the equations. This constant, later known as the *cosmological constant*, is one of the simplest models of dark energy, currently in agreement with a wide variety of cosmological observations.

The first indirect evidences for dark energy come from measurements of galaxy clustering in the 1980's (Peebles 1984; Maddox et al. 1990; Efstathiou et al. 1990). At the end of the 1990's, two independent teams measured the expansion of the Universe using type-Ia supernovae as standard candles (Riess et al. 1998; Perlmutter et al. 1999). Their observations could only be explained by an Universe containing around 30 per cent of matter and 70 per cent of dark energy in the form of a cosmological constant. Few years later, first measurements of the temperature fluctuations in the cosmic microwave background were consistent with an Universe with a flat space (Balbi et al. 2000; de Bernardis et al. 2000), which in combination with measurements of clustering, supernovae and the local expansion rate (Mould et al. 2000), confirmed the acceleration of our Universe's expansion. Age estimates of globular stellar clusters, which are supposed to be among the oldest astrophysical objects, also indicated that the Universe had to be older than the age predicted by models of an Universe only filled with matter (Chaboyer 1998). The following decade, the 2000's, was enriched with the first measurements of the baryon acoustic oscillations in the distribution of galaxies (Eisenstein et al. 2005; Cole et al. 2005). Used as a standard ruler, baryon acoustic oscillations re-confirmed the accelerated nature of the expansion.

Since the 2010's, several cosmological experiments were constructed with the goal of improving the precision of all the measurements discussed above. Today, in the 2020's, state-of-the-art measurements of cosmological parameters achieved relative uncertainties of the order of a few percent or less, a quite remarkable score for observational cosmology. Nevertheless, the data shows no evidence for a departure from a model of Universe governed by GR, composed by 70 per cent of this mysterious cosmological constant and 25 per cent by this mysterious dark matter.

The future is promising with several stage-IV experiments starting their data taking. They will measure millions of galaxy redshifts, billions of galaxy fluxes and shapes, hundreds of thousands of type-Ia supernovae, the cosmic microwave background temperature and polarisation to exquisite precision and resolution, and many more cosmic probes. We truly hope that the data will shed a light over some of the largest mysteries involving our own history.

## 1.3 Model of an expanding Universe

### 1.3.1 Assumptions and ingredients

In order to define some important cosmological parameters, I will quickly review the basics of the currently most accepted model for our Universe in the limit of very large scales, where it can be considered homogeneous and isotropic. This is also known as the *background* model, since the formation of structures (see next section) can be modelled as small perturbations on top of it. It is also assumed that gravity is described by general relativity (GR), which interlinks the fabric of space-time and energy densities of the Universe constituents: *baryons* (protons, electrons, atoms), photons, neutrinos, cold dark matter and dark energy.

The expansion of space is usually parametrised by the scale factor  $a(t)$ , a function of time  $t$  that dictates how physical and comoving distances relate to each other. We assume that  $a(t_0) = 1$  today ( $t_0 = 13.8$  Gyr) and that  $a \rightarrow 0$  as  $t \rightarrow 0$ . The speed of the expansion and its acceleration are simply first and second derivatives with respect to time,  $\dot{a}(t)$  and  $\ddot{a}(t)$ , respectively. The expansion rate of the Universe is defined as  $H(t) = \dot{a}(t)/a(t)$  and takes the Hubble constant value today,  $H(t_0) = H_0 = 100h$  km/s/Mpc  $\sim 70$  km/s/Mpc. Since the scale factor is a monotonically increasing function of time (in our simple model),  $a$  is commonly used to describe a given cosmic epoch.

One of the consequences of the expansion is that relativistic species, such as photons and massless neutrinos, loose energy or are *redshifted* as they propagate. In astronomy, the redshift  $z$  is defined as the relative difference in wavelength between the one emitted by source and the one observed. Thus, in an expanding Universe the scale factor is related to the redshift as  $a = 1/(1+z)$ . Redshift is also commonly used to describe cosmic epoch or distances. The time today  $t_0$  corresponds to  $z = 0$  while  $z \rightarrow \infty$  as  $t \rightarrow 0$ .

### 1.3.2 The expansion rate

The expansion rate can be derived from GR equations and be written as a function of redshift and the densities of each constituent

$$H^2(z) = H_0^2 [\Omega_m(1+z)^3 + \Omega_r(1+z)^4 + \Omega_{de}(z) + \Omega_k(1+z)^2] \quad (1.1)$$

where  $\Omega_x = \rho_x/\rho_{\text{crit}}$  is the ratio between the density constituent  $x$  today and the critical energy density  $\rho_{\text{crit}} = 3H_0^2/8\pi G$ , which is the density required for a flat space geometry. The subscripts  $m$ ,  $r$ ,  $de$  and  $k$  stand for non-relativistic species (baryons, cold dark matter and



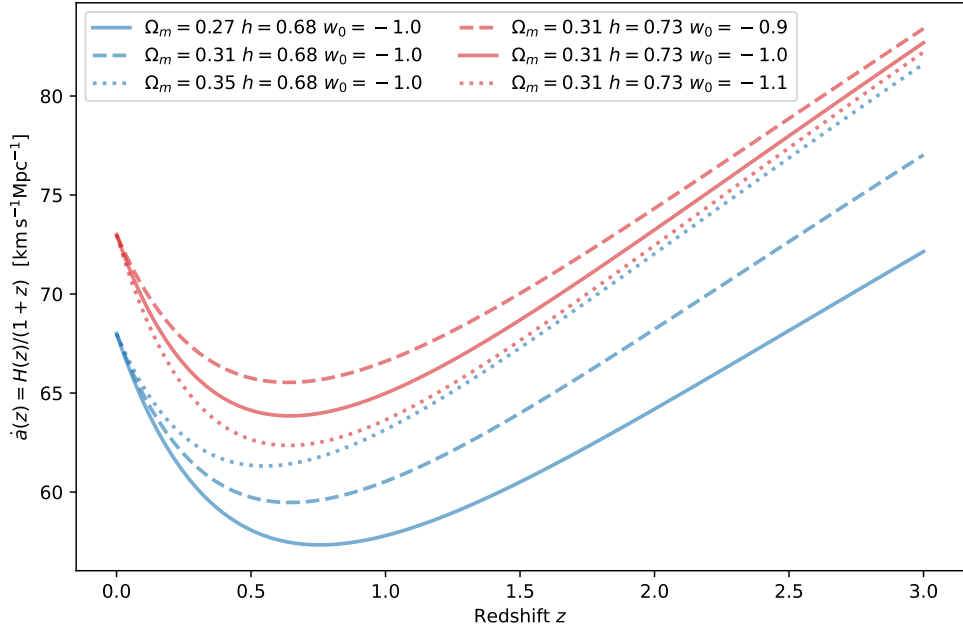


Figure 1.1: The Hubble expansion rate divided by  $(1+z)$ , i.e., the speed of expansion  $\dot{a}(z)$ , versus redshift  $z$  for several cosmological models.

non-relativistic neutrinos), relativistic species (photons and relativistic neutrinos), dark energy and curvature, respectively. The curvature “density” parameter is defined as a function of the others,  $\Omega_k = (1 - \sum_x \Omega_x)$ , and it is zero in a flat space geometry. Non-relativistic species dilute in an expanding universe so their energy density is proportional to  $a^{-3} = (1+z)^3$ . Relativistic species dilute as well but have an extra factor of  $a^{-1}$  due to redshifting. The dark energy term is written as a general function of redshift in Eq. 1.1. In the simplest case of a cosmological constant,  $\Omega_{\text{de}}(z) = \Omega_\Lambda$ , i.e. a constant. The dependency of curvature with  $a^{-2}$  simply comes from the field equations.

Figure 1.1 shows the Hubble expansion rate (times  $a(z)$ ) as a function of redshift  $z$  for a few example of cosmological parameters. All models considered have zero curvature. For all models, the Universe decelerates ( $\dot{a}$  decreases) as time goes by (redshift decreases) until dark energy starts to dominate around  $z \sim 0.5 - 0.7$ , when the expansion accelerates until today. We can see how the Hubble constant  $h$  simply sets the overall amplitude of the curve. Also, for increasing values of  $\Omega_m$ , the transition between matter-dominated to dark energy-dominated era happens at lower redshift.

### 1.3.3 Distances in cosmology

Distances are non-trivially defined in an expanding, potentially non-flat, Universe. Therefore, different types of distances are more appropriate for certain observables. The *Hubble distance* is related to the expansion rate as

$$D_H(z) = \frac{c}{H(z)}, \quad (1.2)$$

and it is commonly used in measurements of baryon acoustic oscillations in the line-of-sight direction (see Chapters 3 and 4) or to represent the size of a causally connected region (in the

future) of the Universe.

The *comoving distance* to an object at redshift  $z$  is written as

$$D_C(z) = \int_0^z dz' D_H(z'). \quad (1.3)$$

This expression yields the distance travelled by a photon from a source towards us but factoring out the scale factor, effectively “removing” the effect of expansion.

The *comoving angular diameter distance*  $D_M(z)$  is useful when distances are inferred from angles, which are affected by the curvature of space. An object with comoving size  $l$  at a redshift  $z$  would be seen with an angular size  $\theta$ . This allows us to define  $D_M(z) = l/\theta(z)$  as

$$D_M(z) = D_C(z) \operatorname{sinc} \left( \sqrt{-\Omega_k} \frac{D_C(z)}{D_H(z=0)} \right), \quad (1.4)$$

where  $\operatorname{sinc}(x) = \sin(x)/x$ . This distance is used in observations of gravitational lensing and in measurements of the baryon acoustic oscillations in the transverse direction to the line-of-sight.

Analogously, the *luminosity distance*  $D_L(z)$  relates the flux  $f$  received by an object at redshift  $z$  with intrinsic luminosity  $L$ . We define then  $D_L(z) = \sqrt{L/4\pi f(z)}$  as

$$D_L(z) = D_M(z)(1+z). \quad (1.5)$$

This distance is used when considering standard candles, where fluxes are used to infer distances. Fluxes of these standard candles vary by several orders of magnitude, so it is handy to define the *distance modulus* as the logarithm of the luminosity distance in units of 10 parsecs (pc):

$$\mu(z) = 5 \log_{10} \frac{D_L(z)}{10 \text{ pc}}. \quad (1.6)$$

This definition is such that it easily relates the observed  $m(z)$  and absolute  $M$  magnitude of an object via  $\mu(z) = m(z) - M$ .

In all distances defined above, the dependency to the cosmological energy densities  $\Omega_x$  happens through  $H(z)$  (Eq. 1.1). Distance measurements allow us to constrain these parameters. However, all cosmological distance measurements are *relative*, which means they are defined to an arbitrary normalisation<sup>2</sup>, such as the size of the standard ruler or the intrinsic luminosity of a standard candle. One can ignore this normalisation and still constrain cosmological densities by performing several distance measurements as a function of redshift. Or one can use some calibrated estimate of this normalisation, and estimate parameters using a single distance measurement. Note that in all the distances defined above, the dependency with  $H_0$  only impacts this arbitrary normalisation, so it cannot be measured without assuming the size of the ruler or the luminosity of the candles. It is useful to factor out this dependency with  $H_0$ , and write distances in units of  $h^{-1}$ Mpc (numerically equivalent to set  $H_0 = 100 \text{ km/s/Mpc}$ ). These are the units used in the vast majority of analyses of the large-scale structures and they will be also used throughout this manuscript.

---

<sup>2</sup>Two exceptions are parallax distances which are an absolute measurements based on the known distance between the Earth and the Sun and gravitational wave distances, which amplitudes depend on the distance and on the masses of the progenitors, though the masses affect the wave form as well, breaking the degeneracy.

### 1.3.4 Dark energy models

In the Eq. 1.1 for the expansion rate of the Universe, the dark energy density term is written as a general function of redshift. As mentioned before, the simplest model is to consider a cosmological constant  $\Omega_{\text{de}}(z) = \Omega_{\Lambda}$ . The cosmological constant can be thought as equivalent to a fluid with negative relativistic pressure  $p_{\text{de}} = -\rho_{\text{de}}$ . This model can be extended by considering a different equation of state  $p_{\text{de}} = p_{\text{de}}(\rho_{\text{de}}) = w(z)\rho_{\text{de}}$ , where  $w(z)$  can be a constant or a more general function of redshift. The only constraint is that  $w(z) < -1/3$ , which is required to obtain an accelerated expansion in a dark energy dominated Universe. One widely known parametrisation is given by  $w(a) = w_0 + w_a(1 - a)$  (Chevallier and Polarski 2001; Linder 2003).

The literature contains a huge variety of models that attempt to be more physically motivated than those just presented (see Weinberg et al. 2013, section 2.2 for a review). Some of them suggest that maybe general relativity breaks down on large enough scales, causing the expansion to accelerate. They suggest extensions or modifications to Einstein's theory of general relativity, and are commonly known as *modified gravity* models.

## 1.4 Model of the large-scale structures

The Universe is clearly not homogeneous and isotropic. Matter clustered under the influence of gravity, creating the cosmic web, composed of clusters, filaments, empty regions. The model described in the previous section only describes the Universe as a whole, without any inhomogeneities. If we want to describe the evolution of the structures we need to consider an extension to the background model.

Observations of the large-scale structures and the cosmic microwave background (CMB) indicate that the matter density field today is the evolution of tiny density perturbations from the early Universe. Temperature fluctuations in the CMB are of the order of  $10^{-5}$ , while matter densities today can reach values several orders of magnitude larger than the average density. The most widely accepted theory is that the CMB fluctuations, and today's structures, are originated from the inflation-grown quantum fluctuations, that evolved under the influence of gravity and pressure in the hot dense plasma-like epoch of our Universe.

It is been possible to model the physics of the evolution of these tiny fluctuations using *linear perturbation theory*. The predictions of this theory are an excellent match to observations of the CMB and to the late-time distribution of matter on large scales, i.e., larger than few tens of Mpc. On smaller scales however, gravity becomes highly non-linear and more advanced calculations are required. There are two main approaches to model non-linearities: theoretical calculations going beyond linear terms in perturbation theory or numerical n-body simulations.

The main idea of perturbation theory is to consider that the density of a given species  $x$  is given by:

$$\rho_x(\vec{x}, t) = \bar{\rho}_x(t) [1 + \delta_x(\vec{x}, t)], \quad (1.7)$$

where the bar indicates average over the whole Universe and  $\delta_x \ll 1$  is a small perturbation. The average density  $\bar{\rho}_x$  follows the background evolution described in section 1.3. A new set of equations can be derived for the evolution of  $\delta_x$  for each species depending if  $x$  denotes relativistic species (photons, hot baryons, hot neutrinos) or non-relativistic species (cold baryons, cold neutrinos or dark matter). Only some families of models also consider fluctuations in the dark energy density, but most commonly dark energy only acts in the background expansion.

The evolution of  $\delta_x$  for a given species is dictated by both Boltzmann and Einstein's GR equations. Boltzmann equations describe the evolution of phase-space distributions of each constituent, considering collisions (pressure) and particle creation and annihilation. GR equations describe how each species influence gravitational potentials which in turn make perturbations grow.

Another important ingredient in these equations is the *peculiar velocities*  $\vec{v}_x(\vec{x}, t)$  of each species, particularly the non-relativistic ones since relativistic species are assumed to have  $v \sim c$ . Peculiar velocities directly impact the evolution of perturbations through Boltzmann, Euler and continuity equations. These velocities can be measured from galaxies if their distances are estimated. They are an important probe of the cosmological model and on the strength of gravity itself. They will be a key actor in chapters 4 and 5, where I will discuss redshift-space distortions and the clustering of velocities.

### 1.4.1 Statistical description of perturbations

The value of the matter density field at a given position  $\vec{x}$  is hardly observable since most of it is in the form of dark matter. Only baryons which emit photons are observable. Furthermore, it is impossible to learn about the physics of a single value of  $\delta_m(\vec{x})$  since we do not have access to its time evolution: our sky seems static on time-scales of cosmological evolution.

In order to connect our physical model of structure growth to the observed density field, we turn our interest to its statistical properties. We believe that density perturbations are the evolved form of quantum initial perturbations, which would follow Gaussian statistics. If Gaussian, all the information of initial perturbations would be contained in its first two moments. The evolution of these perturbations through time modifies these initial statistical properties, which become highly complex, for instance creating non-Gaussianities. Therefore, the final density field is described by all higher order moments as well.

If we focus on the density field on large scales, most of its cosmological information is still contained in the first moments, starting from the two-point function (the one-point function is zero by definition), which is defined as

$$\xi(\vec{x}, \vec{y}) = \langle \delta(\vec{x})\delta(\vec{y}) \rangle \quad (1.8)$$

where  $\langle \cdot \rangle$  denotes an ensemble average. Given that we only have one single realisation of the Universe, we assume that averages over space (in an infinite space) are equivalent to ensemble averages.

The assumptions of homogeneity and isotropy of the background can be applied in a statistical sense to the perturbations. Statistical homogeneity makes  $\xi$  not to depend on the specific locations  $\vec{x}$  and  $\vec{y}$  but only on their separation  $\vec{r} = \vec{y} - \vec{x}$ . Statistical isotropic makes  $\xi$  no longer depend on the orientation of  $\vec{r}$ , only on its absolute value  $r = |\vec{r}|$ . Therefore, the two-point correlation function simplifies to a function of  $r$ :

$$\xi(r) = \langle \delta(\vec{x})\delta(\vec{x} + \vec{r}) \rangle. \quad (1.9)$$

The correlation function  $\xi(r)$  is a powerful observable which can be compared to predictions from a given cosmological model. Most of my work was dedicated to measuring correlations with galaxy survey data.

### 1.4.2 Configuration and Fourier space

The time evolution of linear density perturbations  $\delta(\vec{x})$  is dictated by a set of second-order differential equations, containing derivatives with respect to time and space. It is thus con-

venient to apply Fourier transforms to these equations, where derivatives with respect to space become simple products. Separations  $\vec{r}$  in the so-called *Configuration space* become wavevectors  $\vec{k}$  in *Fourier space*. The density perturbations in Fourier space are defined by

$$\tilde{\delta}(\vec{k}) = \int d^3x e^{-i\vec{k}\cdot\vec{x}} \delta(\vec{x}), \quad (1.10)$$

and its inverse is

$$\delta(\vec{x}) = \frac{1}{(2\pi)^3} \int d^3k e^{+i\vec{k}\cdot\vec{x}} \tilde{\delta}(\vec{k}), \quad (1.11)$$

where integrals are assumed to be performed over an infinite volume.

As previously, we are interested in the two (or more) point functions of the field. In Fourier space, the two-point correlation function is known as the *power spectrum* and is defined as

$$\langle \tilde{\delta}(\vec{k}) \tilde{\delta}(\vec{k}') \rangle = (2\pi)^3 P(k) \delta_{\text{Dirac}}^3(\vec{k} - \vec{k}'), \quad (1.12)$$

where  $\delta_{\text{Dirac}}^3(\vec{x})$  is a three-dimensional Dirac distribution. The assumption of spatial homogeneity is guaranteed by the properties of Fourier space, while the assumption of isotropy makes the power spectrum to be only a function of the absolute value of the wavevector  $k = |\vec{k}|$ .

When considering the evolution of perturbations in Fourier space to the linear level, each mode  $\tilde{\delta}(\vec{k})$  evolves independently of the others. We often say that there is no *mode coupling* in linear theory.

The power spectrum  $P(k)$  is an observable, as powerful as the correlation function, to constrain cosmological models. While in theory these two contain exactly the same compressed statistical information, in practice the analyses of real data use different finite ranges of scales which are affected differently by systematic effects. Therefore, analyses of real data in configuration and Fourier space complement each other. My contributions to such analyses are presented in chapter 4.

### 1.4.3 Cosmological dependency of the power spectrum

The power spectrum  $P(k)$  or the correlation function  $\xi(r)$  of density perturbations are excellent probes of the cosmological model. They encode information accumulated over the whole Universe's history, since the Big Bang until today.

At the end of inflation, we think that perturbations were roughly Gaussian (as are quantum fluctuations of the vacuum), having a nearly scale independent power spectrum defined by

$$P_0(k) = A_s k^{n_s-1}, \quad (1.13)$$

where  $A_s \sim 2 \times 10^{-9}$  is the amplitude of scalar perturbations and  $n_s \sim 0.96$  is the scalar spectral index. The term scalar refers to standard density (or gravitational potential) perturbations, while tensorial perturbations (of the space-time metric) refer to primordial gravitational waves. Vectorial perturbations decay and rapidly become negligible in most common cosmological models.

After inflation, the majority of the energetic budget of Universe was held by relativistic species: photons and neutrinos. Since their energy density decays as  $a^{-4}$ , their contribution quickly drops. This radiation-dominated era lasted until  $z \sim 5000$ , or  $t \sim 20$  kyr, when the energy density of matter became the dominant source. Matter density decays at a slower rate, proportional to  $a^{-3}$ . dark energy became dominant at  $z \sim 0.5$ , or  $t \sim 9$  Gyr, since its

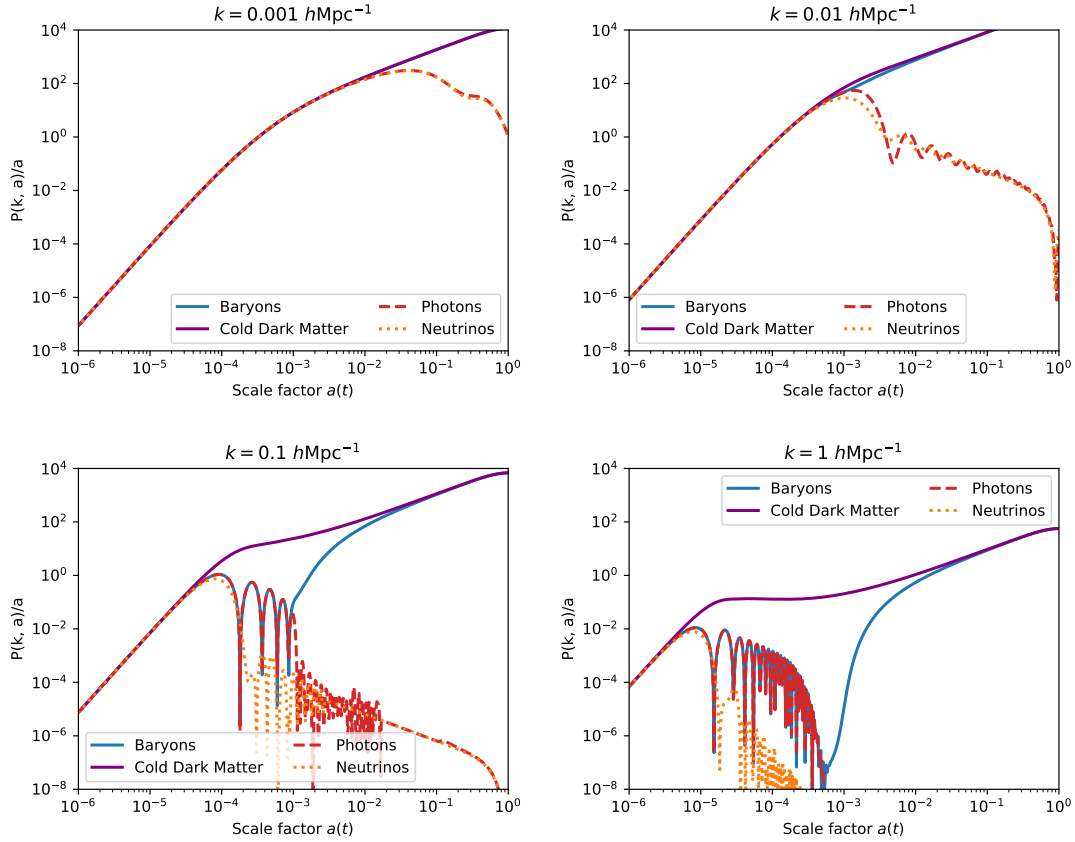


Figure 1.2: Evolution of the amplitude of the power spectrum versus the scale factor for four different species and four different scales (one per panel). The evolution was computed using CAMB. A movie version can be found at [https://github.com/julianbautista/movie\\_correlations](https://github.com/julianbautista/movie_correlations).

energy density is constant (or close to constant). The actual duration of each era depends on relative values of the energy densities, parametrised by  $\Omega_x$  (see 1.3.2).

Matter perturbations grew at different rates during these different eras. Figure 1.2 shows how the power spectrum grows over time for different species and for different scales. In radiation-dominated era, large-scale modes (large values of  $r$  or small values of  $k$ ) grew while small-scale ones oscillated, due to battle between gravity and radiative pressure. These are known as *baryon acoustic oscillations*. These oscillations are imprinted in the power spectrum of the photon-baryon fluid, at the time of the CMB, and in the matter power spectrum at later times. The radiative pressure prohibited the growth of small-scale perturbations, so the matter power spectrum is damped, which results in this hill-like shape. In the matter-dominated era, the radiative pressure becomes negligible and all modes grow equally (neglecting non-linear growth). When dark energy is dominant, the expansion accelerates and growth of structures slows down slightly.

Well after recombination, matter (baryons + dark) density perturbations obey the following differential equation, if we assume GR and linear theory:

$$\ddot{\delta}(\vec{x}, t) + 2H(z)\dot{\delta}(\vec{x}, t) - \frac{3}{2}\Omega_m H_0^2(1+z)^3\delta(\vec{x}, t) = 0. \quad (1.14)$$

If we assume that  $\delta(\vec{x}, t) = \delta(\vec{x}, t_0)G(t)/G(t_0)$ , we can factor out the spatial dependency and only solve for the time-dependent term  $G(t)$ , known as the *growth factor*. Its logarithmic derivative with respect to the scale factor  $a(t)$  is called *growth rate* of structures:

$$f(a) \equiv \frac{d \ln G(a)}{d \ln a}, \quad (1.15)$$

which is a key observable for which the predicted value depends on the cosmological parameters and on the assumed theory of gravity (e.g., GR in Eq. 1.14).

In linear theory, the peculiar velocity field is simply related to density perturbations via the continuity equation

$$\nabla \cdot \vec{v}(\vec{x}, a) \equiv \theta(\vec{x}, a) = -a^2 H(a) \frac{d\delta(\vec{x}, a)}{da} = -aH(a)f(a)\delta(\vec{x}, a) \quad (1.16)$$

where  $\theta$  is the velocity divergence, a convenient scalar field describing velocities. In Fourier space, the velocity field is simply proportional to the density. The velocity divergence power spectrum  $P_{\theta\theta}$  is an important ingredient for models of redshift-space distortions, as we discuss in the following section.

#### 1.4.4 The amplitude of the power spectrum

The matter power spectrum  $P(k)$  represents the variance of density perturbations at a given wavenumber  $k$ . The total variance of the density field is therefore the sum of all contributions over all available three dimensional modes. This is equivalent to the correlation function at zero separation:

$$\sigma^2 = \xi(r=0) = \frac{1}{(2\pi)^3} \int d^3k P(k) = \frac{1}{2\pi^2} \int_0^\infty dk k^2 P(k), \quad (1.17)$$

which diverges if we use the linear prediction for the matter power spectrum. This is because the linear power spectrum is the evolution of a nearly scale invariant power spectrum, equivalent to white noise, so the variance simply keeps increasing as we consider larger ranges of scales. Therefore, it is convenient to smooth the linear density field using a three-dimensional top-hat filter with radius  $R$ . The variance of the smoothed field is then

$$\sigma_R^2 = \frac{1}{2\pi^2} \int_0^\infty dk k^2 P(k) W_R^2(k), \quad (1.18)$$

where  $W_R(k) = 3[\sin(kR) - kR \cos(kR)]/(kR)^3$  is the Fourier transform of the top-hat filter in three dimensions.

In studies of large-scales structures, it is common to use  $\sigma_R^2$  as a parameter defining the amplitude of the linear power spectrum. As discussed before, this amplitude depends on the values of  $A_s$  (Eq. 1.13) and  $h$ , which determines the age of the Universe and how long structures could have grown. Therefore  $\sigma_R^2$  is degenerate with  $A_s$  and  $h$  though is more representative of the amplitude of  $P(k)$  at later times. Other energy density parameters  $\Omega_x$  also affect this amplitude, but these also change the overall shape and are not degenerate with  $\sigma_R$ .

The variance  $\sigma_R$  is a decreasing function of  $R$ . Historically, the chosen values for  $R$  yield variances near unity, which correspond to the regime where linear theory should break. Recently, the value of  $R = 8h^{-1}\text{Mpc}$  has been used in several analysis, though it has been argued that choosing  $R = 12\text{Mpc}$  (without the  $h$  dependency) is a better choice to break degeneracies (Sánchez 2020). The accepted value for  $\sigma_8$  is around 0.8.

In measurements of redshift-space distortions and peculiar velocities, there is a degeneracy between the value of  $f$  and the amplitude of the power spectrum,  $\sigma_8$ . Therefore, measurements can only constrain the combination  $f\sigma_8$  (see more in chapter 4).

## 1.5 Cosmological probes of expansion

In this section, I present the basics of the cosmological observables that allow us to learn about dark energy. Some of these are mostly sensitive to the background evolution presented in section 1.3.2, while others depend on the matter perturbations and their statistical properties as discussed in section 1.4.

### 1.5.1 Direct measurements of $H_0$

The Hubble constant  $H_0$  is probably one of the first cosmological parameters to be estimated by the person giving its name to it (Hubble 1929). Today, there are a few techniques to estimate  $H_0$  which yield roughly independent results. See Riess 2020 for a quick review of the latest results.

The most traditional method is known as the *distance ladder*. The idea is to measure the distance-redshift relationship of objects well in the Hubble flow, i.e., far enough so they feel the expansion. The distance to the closest objects in the Solar neighbourhood can be estimated with the parallax method. The Gaia satellite has the current largest catalogue of parallax measurements, containing more than a billion stars. Parallax is one of the few direct distance estimating methods. They can be used to measure the intrinsic luminosity of some objects thought to be standard candles, such as Cepheid stars, RR Lyrae, and the largest red giant stars. One can use these candles to estimate distances to even more distant objects, explaining the usage of the word *ladder*. Riess et al. 2021 contains the latest measurements using Gaia parallaxes, Cepheids and type-Ia supernovae to determine  $H_0 = 73.0 \pm 1.4$  km/s/Mpc. Freedman et al. 2019 is latest measurement of distances using the tip of the red giant branch.

The distance to the galaxy NGC 4258 could be determined thanks to the presence of a water maser orbiting the center of this galaxy. The proper motions of several clouds orbiting close to the central massive black hole could be measured both in radial and angular directions, strongly constraining the dynamics of the system (Herrnstein et al. 1999). The most recent measurement yields  $D = 7.576 \pm 0.082$  (stat.)  $\pm 0.076$  (sys.) Mpc (Reid et al. 2019), which can be used as an alternative to Cepheids to anchor the distance ladder and provide an estimate to  $H_0$ .

Another alternative method is based on the time-delay of the signals emitted from a quasar behind a strong gravitational lens. The lens creates multiple images of the same background quasar, but the light paths have slightly different lengths. Since quasars are variable objects, the same variability is observed with a delay of several days between the different images. By modelling the distribution of matter between the quasar and us, it is possible to convert these time delays into an estimate of  $H_0$ . The collaboration named “ $H_0$  Lenses in COSMOGRAIL’s Wellspring”, or H0LICOW, produced the latest comprehensive measurement of  $H_0$  using strongly-lensed quasars (Wong et al. 2019). An alternative measurement has been performed using strong lenses from the Sloan (SLACS). They are compared to H0LICOW in Birrer et al. 2020.

More recently, gravitational waves from a merger of two black holes were observed for the first time by the LIGO collaboration (T. L. S. Collaboration and the Virgo Collaboration 2016). Not long after, a merger of two neutron stars was observed by both LIGO and Virgo, but this time an electromagnetic counterpart was also detected (T. L. S. Collaboration and T. V. Collaboration 2017), pointing to the galaxy where the event occurred. Given the well predicted shape of the wave form and its dependency on the masses of the neutron stars, it was possible to estimate the luminosity distance to the host galaxy independently of its



redshift. By combining this distance with an actual measurement of the host galaxy redshift, a single event could provide a rough estimate of  $H_0$  (Abbott et al. 2017). This measurement opened the field of cosmology using standard sirens, i.e., gravitational wave sources.

The cosmic microwave background can yield an estimate of  $H_0$  but it is strongly degenerate with other unknown parameters, such as curvature or dark energy densities. Some of these degeneracies are reduced when considering the effect of gravitational lensing of the CMB or when combining with other probes of late times. Therefore, precise measurements of  $H_0$  from the CMB alone are only possible when considering more restrictive models, such as a flat space with a cosmological constant.

### 1.5.2 Type-Ia supernovae

As mentioned earlier, type-Ia supernovae (SNIa) can be used as standard candles for absolute distance measurements, if they are anchored by another method. The SNIa can also be used without anchors to produce relative distance measurements and constrain dark energy. In this case, the only assumption is that the SNIa intrinsic brightness does not evolve with redshift.

The main observable are the fluxes of SNIa as a function of time in different photometric bands, known as light-curves. Spectroscopic follow up observations of the explosion can confirm the type of the supernova based on the features present in their spectra. Another key ingredient is the redshift of the host galaxy of the SNIa, which is commonly measured with spectroscopy as well.

Spectro-photometric models of SNIa are used to fit the observed light-curves, yielding their apparent magnitudes at peak luminosity in a given photometric band. These models account for correlations between colour and duration of light curves and the peak magnitude, reducing the intrinsic scatter in these magnitudes from 40% to roughly 15%. It is essential to obtain accurate and precise measurements of SNIa fluxes in order to obtain the best cosmological constraints. Large sets of realistic simulations of the data are required in order to correct for selection effects. The final product of the analysis is a set of distance moduli (Eq. 1.6) and their host-galaxy redshift, which can then be compared to models of expansion of the Universe.

Distance moduli depend on dark energy through the integral of  $H^{-1}(z)$  (Eq. 1.5 and 1.6). In order to obtain good constraints on dark energy properties, it is important to have a large redshift coverage, at least covering the transition between matter-dominated to dark energy-dominated eras, so over  $0 < z < 1$ . The advantage of SNIa for dark energy studies is that they can span these redshifts with a high sampling rate, which helps in the study of the expansion rate. One of the inconveniences is that SNIa are complex and poorly understood astrophysical events, the intrinsic scatter in luminosity cannot be reduced to better than 12%, potentially limiting the gains constraining power from future experiments.

The latest comprehensive study of SNIa combines data from more than a dozen projects into a single sample, the Pantheon sample, from which dark energy constraints were derived (Scolnic et al. 2018). The Dark Energy Survey also measured a more recent sample of few hundreds of SNe (Brout et al. 2019b; Brout et al. 2019a; Kessler et al. 2019), but not all have spectroscopic classification. Their constraints on dark energy are not yet competitive compared to the Pantheon sample. The Zwicky Transient Facility (ZTF, Graham et al. 2019) is currently observing the Northern sky on a search for transient events and is expected to discover around 5000 spectroscopically confirmed SNIa at low redshifts ( $0 < z < 0.12$ ) until 2023. After 2023, the Rubin Observatory Legacy Survey of Space and Time (LSST) will take over in

the Southern hemisphere and will discover more than 300 thousand SNIa up to  $z < 0.5$ .

While ZTF and Rubin's samples of SNIa will not decrease significantly the errors on dark energy parameters, they cover a redshift range where other powerful probes, such as baryon acoustic oscillations or weak-lensing, lack of statistical power due to limited volume. Furthermore, at lower redshifts when dark energy is dominant, SNIa are complementary to redshift-space distortions (RSD) when it comes to testing the validity of GR or constraining alternative models of gravity, as solutions for dark energy.

Not only SNIa are a great probe of the expansion history, they can also provide peculiar velocities of their host galaxies via their inferred distances. These peculiar velocities and their statistical properties can complement RSD analysis when measuring the growth-rate of structures (Eq. 1.15). Estimates of  $H(z)$  and  $f(z)$  can help break degeneracies between simple dark energy models and more involved models of gravity (Kim and Linder 2020; Graziani et al. 2020). Chapter 5 is dedicated to this topic, to which I plan to dedicate the next few years of my research.

### 1.5.3 Big Bang nucleosynthesis

In the post-inflation Universe, when temperatures are below the equivalent of 100 MeV, perturbations in the matter and radiation fields are quite small and most of the physics is dictated by the interactions between protons, electrons, neutrons, and photons, as described by the Boltzmann equations. As the Universe cools down and rarefies, protons and neutrons start forming atoms of deuterium, tritium, helium and heavier elements. This process is known as the Big Bang nucleosynthesis (BBN).

The relative amount of each of the formed elements depends on the expansion rate of the Universe at that time as well as the physical density of baryons  $\omega_b = \Omega_b h^2$  and radiation  $\omega_r = \Omega_r h^2$ . Given that the energies involved are within reach of particle accelerators, these reactions can be studied with great detail on Earth, allowing us to build accurate models of the BBN (see Pitrou et al. 2021 for the latest calculations and references therein).

Observations of the primordial abundances can be compared to the predictions by BBN models. Abundances of deuterium, helium, and others can be estimated from spectroscopic observations of HII regions in metal-poor galaxies or from absorption lines of the intergalactic medium in quasar spectra. The most up-to-date measurements of the primordial helium-4 abundance yields  $Y_p = 0.2453 \pm 0.0034$  (Aver et al. 2021), while the deuterium one is  $D/H = (2.527 \pm 0.030) \times 10^{-5}$  (Pitrou et al. 2021). While observations are consistent with BBN models for deuterium and helium, the abundance of lithium-7 exhibits a factor 3 discrepancy, which is a huge problem in BBN but quite often neglected.

The temperature fluctuations in the cosmic microwave background (CMB) are also very sensitive to the physical baryon density  $\omega_b$ . Historically, the values obtained from CMB have been in good agreement with BBN measurements, showing that baryons make up to around 16% of the total matter content of the Universe, or  $\omega_b = (2.195 \pm 0.022) \times 10^{-2}$ . The agreement between two quite independent probes is one of the great successes of the current cosmological model, though they also enforce the need for a dark matter component.

### 1.5.4 Baryon acoustic oscillations

Baryon acoustic oscillations (BAO) is the name given to the propagation of sound waves in the primordial plasma (baryons and photons), prior to recombination. Because of the high pressure on small scales at those times, each initial density perturbation had a spherical density wave around them propagating outwards at the speed of sound in that medium. The

speed of sound in the plasma is given by

$$c_s(z) \equiv \sqrt{\frac{1}{3[1+R(z)]}} \quad (1.19)$$

where  $R(z) = 3\rho_b/4\rho_\gamma$  is the baryon-to-photon ratio. The propagation of sound waves occurred until the temperatures and densities dropped to values such that baryons no longer felt the pressure from photons, known as the *drag epoch*, which is close in time to the recombination (but technically not the same epoch), at  $z \sim 1100$ . This process left a slight overdense shell of matter around each initial perturbation with a radius given by

$$r_{\text{drag}} \equiv r_s(z_{\text{drag}}) = \int_{\infty}^{z_{\text{drag}}} dz' \frac{c_s(z')}{H(z')} \quad (1.20)$$

which is known as the *sound horizon at drag epoch*, or the BAO scale. Today,  $r_d$  has a physical size of about 147 Mpc, much larger than any collapsed structure in the Universe. As one can see from Eqs. 1.19 and 1.20,  $r_{\text{drag}}$  mainly depends on  $\omega_b$  and  $\omega_c$  assuming the CMB gives a precise estimate of  $\omega_\gamma$ . The dependency of  $r_{\text{drag}}$  is mainly through  $\omega_b$ .

After recombination, the sound horizon scale only increases in size due to the expansion, or equivalently, its comoving size remains unchanged. Therefore, the BAO scale is a great *standard ruler* to study the expansion rate of the Universe. In practice, the BAO scale is observed statistically in the two-point function of the matter density field, so it is often classified as an statistical standard ruler. In configuration space, the correlation function  $\xi(r)$  presents a small peak at separations corresponding to the BAO scale  $r_d$ , while in Fourier space the power spectrum  $P(k)$  contains an oscillatory pattern as a function of scale with a period proportional to the BAO scale. In chapters 3 and 4, I present my past work in the measurement of the BAO scale using  $\text{Ly}\alpha$  forests and galaxies, respectively, as tracers of the matter density field.

Given that a galaxy survey is made of angular positions and redshifts, which are true observables, the BAO peak is effectively measured as an angle  $\Delta\theta_{\text{BAO}}$  or as a difference in redshift  $\Delta z_{\text{BAO}}$ . These can be modelled as ratios of distances to the BAO scale  $r_d$  as

$$\Delta\theta_{\text{BAO}}(z_{\text{eff}}) = \frac{D_M(z_{\text{eff}})}{r_d} \quad (1.21)$$

$$\Delta z_{\text{BAO}}(z_{\text{eff}}) = \frac{D_H(z_{\text{eff}})}{r_d} \quad (1.22)$$

where  $z_{\text{eff}}$  is the effective redshift of the galaxy survey,  $D_M$  is the comoving angular diameter distance (Eq. 1.4), and  $D_H$  is the Hubble distance (Eq. 1.2).

There are two ways BAO can be used to constrain cosmological models, depending if we assume that  $r_{\text{drag}}$  is known or not. If  $r_{\text{drag}}$  is known, i.e. given by Eq. 1.20 using some value for  $\omega_b$  (given by the CMB or BBN for instance), then BAO measurements are converted to absolute distance measurements which depend only on  $H(z)$ . Therefore, BAO can constrain  $H_0$ , curvature and dark energy. If  $r_{\text{drag}}$  is supposed to be unknown but still a standard ruler, then BAO constraints the ratio of  $E(z) = H(z)/H_0$  to the combination  $H_0 r_{\text{drag}}$  which is now degenerate. This is similar to the case of type Ia supernovae, where their absolute magnitude is degenerate with  $H_0$ . By combining several BAO measurements at different effective redshifts, BAO is a powerful probe of dark energy and curvature.

Current BAO measurements span effective redshifts from 0.1 to 2.3. The Sloan Digital Sky Survey (SDSS, Eisenstein et al. 2011; Blanton et al. 2017) has measured more than 2 million

redshifts spectroscopically in the past twenty years, producing the largest maps to date of the distribution of matter in the Universe. In addition to SDSS, surveys such as FastSound (Okumura et al. 2016), Vipers (Pezzotta et al. 2017), 6 degree field galaxy survey (6dFGS, Beutler et al. 2012) and WiggleZ (Parkinson et al. 2012) also produced BAO measurements using galaxies, though the volumes probed and the number of galaxies is inferior to SDSS. The latest cosmological constraints from BAO are described in Alam et al. 2021.

### 1.5.5 The cosmic microwave background

The cosmic microwave background (CMB) is one of the richest cosmological probes of all. The photons we receive today last scattered on baryons at a redshift of about 1100, corresponding to roughly 380 000 years after the Big Bang. It is currently the oldest information that we can measure from the Universe today.

The average temperature  $T_{\text{CMB}} = 2.72548 \pm 0.00057$  K (Mather et al. 1994; Fixsen 2009) of these photons, measured by the COBE satellite, tell us on how much energy density from radiation there is in the Universe. Considering a black body Bose-Einstein distribution for the photons, we have that

$$\rho_\gamma = \frac{\pi^2 k^4}{15 \hbar^3 c^3} T_{\text{CMB}}^4 \quad (1.23)$$

The fluctuations around this average temperature,  $\Delta T$ , and the polarisation of these photons trace the structures back at the recombination epoch. Given its early-times nature, the temperature and polarisation fields are extremely well described by Gaussian statistics. Most of the information is therefore contained in the two point functions of these fields. Given that these fields are functions of the position angle in the sky  $(\theta, \phi)$ , it is convenient to decompose these fields into a basis of spherical harmonics  $Y_{\ell m}(\theta, \phi)$ . The amplitude of each harmonic is denoted  $a_{\ell m}$ . The angular power spectrum is simply the variance of these amplitudes:  $C_\ell = \langle a_{\ell m} a_{\ell m}^* \rangle$ , which in linear theory is simply a function of  $\ell$  and independent of  $m$ . The temperature and both E and B polarisation modes of the CMB can be decomposed into spherical harmonics, so we can estimate all cross power spectra, e.g.,  $C_\ell^{TE} = \langle a_{\ell m}^T a_{\ell m}^{E*} \rangle$  is the cross temperature and E-mode polarisation power spectrum.

The temperature and polarisation auto and cross power spectra are exquisitely well modelled by linear perturbation theory. In the most basic  $\Lambda$ CDM model, the power spectrum is a function of only six parameters: the angular scale of baryon acoustic oscillations  $\theta_*$ , the physical density of baryons  $\omega_b$  and dark matter  $\omega_c$ , the primordial power spectrum amplitude  $A_s$  and slope  $n_s$ , and the optical depth to the last scattering surface  $\tau$ . The fact that such a model with so few free parameters can describe so well CMB observations is one of the greatest achievements in modern physics. A great description of the CMB physics from first principles can be found in Dodelson and Schmidt 2020. Several codes are available to compute models for the CMB such as CAMB<sup>3</sup> (Lewis et al. 2000) and CLASS<sup>4</sup> (Lesgourgues 2011).

Since COBE measurements, all experiments focus on measuring exclusively anisotropies in temperature and polarisation. The current reference results based on the CMB come from the Planck satellite. (Planck Collaboration et al. 2020). Since it is a satellite, it has access to the full sky and thus the very large scale modes. The angular resolution of the instrument of about 5 arcmin allows a precise measurement of the power spectrum up to  $\ell \sim 2500$ . Eight acoustic peaks are observed, yielding an extremely tight constraint on  $\theta_*$ . The relative amplitudes of

<sup>3</sup><https://camb.info/>

<sup>4</sup>[https://lesgourg.github.io/class\\_public/class.html](https://lesgourg.github.io/class_public/class.html)

the acoustic peaks is sensitive to  $\omega_b$  and  $\omega_c$ . The overall amplitude of the spectrum yields the  $A_s e^{-2\tau}$ , while its dependency with scale yields  $n_s$ . Polarisation spectra break the degeneracy between  $A_s$  and  $\tau$ . A weak constraint on dark energy is possible thanks to the integrated Sachs-Wolfe effect at low redshift.

On Earth, several experiments were build achieving higher angular resolution or a higher depth in observations. Both the Atacama Cosmology Telescope (ACT, Aiola et al. 2020) and the South Pole Telescope (SPT, Aylor et al. 2017; Balkenhol et al. 2021; Dutcher et al. 2021), measured temperature and polarisation power spectra on a smaller fraction of sky but reaching up to  $\ell \sim 9000$ . The BICEP/Keck (BICEP/Keck et al. 2021) focused on the polarisation signal over large angular scales (low  $\ell$ ) with the goal of measuring the effect of primordial gravitational waves caused by inflation. Due to strong contamination by polarised emission from Galactic dust, currently there is no detection of such primordial signal. New experiments will increase even further the sensitivity and the number of detectors, in order to achieve this goal.

### 1.5.6 Redshift-space distortions

Matter is not static in the Universe. Due to gravity, matter flows from underdense regions towards overdense ones, so it can be described as a fluid following some velocity field. These velocities are commonly called *peculiar velocities*. The radial component of peculiar velocities with respect to the observer alters the observed redshift of an object, due to the Doppler effect. The total observed redshift  $z_{\text{obs}}$  is a combination of its cosmological redshift  $z_{\text{cos}}$ , due to the expansion of the Universe, and the peculiar redshift  $z_{\text{pec}}$ , due to the Doppler effect:

$$1 + z_{\text{obs}} = (1 + z_{\text{cos}})(1 + z_{\text{pec}}) \quad (1.24)$$

where

$$z_{\text{pec}} = \sqrt{\frac{1 + \frac{\vec{v} \cdot \hat{n}}{c}}{1 - \frac{\vec{v} \cdot \hat{n}}{c}}} - 1 \approx \frac{\vec{v} \cdot \hat{n}}{c} \quad (1.25)$$

The first equality is the definition of the relativistic Doppler effect,  $\vec{v}$  is the velocity and  $\hat{n}$  is the radial unitary vector. The right-hand side is the non-relativistic approximation, which is accurate for typical velocities of matter flows in our Universe, of the order of few hundreds of km/s.

When converting observed redshifts into comoving distances, peculiar velocities introduce a small error such that

$$s(z_{\text{obs}}) \approx r(z_{\text{cos}}) + \frac{(1 + z_{\text{cos}})}{H(z_{\text{cos}})} \vec{v} \cdot \hat{n}, \quad (1.26)$$

where  $s$  is the *redshift-space* distance and  $r$  is the *real-space* distance. These errors distort the observed distribution of matter/galaxies in the radial direction. This effect is observable and manifests as anisotropy (radial versus transverse) in the two-point functions of the density field. These anisotropies are named *redshift-space distortions* (RSD). The RSD are a powerful probe of the dynamics of the matter field, i.e., how velocities are related to the densities, and how they contribute to the growth of structures over time. More generally, observations of RSD can be used to test the validity of general relativity since it is a probe of the strength of gravity.

The density contrast of matter is modified when observed in redshift-space relative to real-space. Mass conservation implies that

$$[1 + \delta_s(\vec{s})] d^3 s = [1 + \delta(\vec{x})] d^3 x, \quad (1.27)$$

such that in Fourier space we obtain, in the linear regime (Kaiser 1987)

$$\delta_s(\vec{k}) = (1 + f\mu^2)\delta(\vec{k}) \quad (1.28)$$

where  $f$  is the growth-rate of structures defined in Eq. 1.15 and  $\mu = \vec{k} \cdot \hat{n} / k$  is the cosine of the angle between the wavevector and the line of sight. The power spectrum in redshift-space assuming linear perturbations is simply

$$P_s(k, \mu) = (1 + f\mu^2)^2 P_m^{\text{lin}}(k) \quad (1.29)$$

This means that the redshift-space matter power spectrum has radial modes enhanced by a factor of  $(1 + f)^2$  (where  $f \sim 1$ ) relative to modes transverse to the line of sight. This enhancement is observable with galaxy surveys. If the normalisation of the linear matter power spectrum is parametrised by  $\sigma_8$  (Eq. 1.18), then the actual measured quantity is the product  $f\sigma_8$ . Transposed to a given effective redshift  $z_{\text{eff}}$  of a galaxy survey, we need to scale  $\sigma_8$ , which is usually defined at  $z = 0$ , using the growth factor  $G(z)$ , such that the measured quantity is  $f(z)\sigma_8 G(z)$ . Some works use the notation  $\sigma_8(z) \equiv \sigma_8 G(z)$ .

The observable  $f(z)\sigma_8(z)$  is mostly sensitive to the amount of dark matter  $\omega_m$ , which drives the growth of structures, and  $H(z)$ , which damps the growth (see Eq. 1.14). Therefore, measurements of  $f(z)\sigma_8(z)$  versus redshift can be used to constrain dark energy models or alternatives to GR.

The latest most relevant measurements were performed using data from the Sloan Digital Sky Survey (SDSS), including

- the SDSS Main Galaxy Sample Howlett et al. 2015,
- the Baryon Oscillation Spectroscopic Survey (BOSS, Alam et al. 2017),
- the extended BOSS luminous red galaxy sample (eBOSS LRG, Bautista et al. 2021; Gil-Marín et al. 2020),
- the eBOSS emission line galaxy sample (eBOSS ELG, Tamone et al. 2020; de Mattia et al. 2021),
- the eBOSS quasar sample (eBOSS QSO, Hou et al. 2021; Neveux et al. 2020),

In chapter 4 I present my contributions to the measurement of the growth-rate of structures using the eBOSS LRG sample (Bautista et al. 2021). The cosmological implications of SDSS growth-rate measurements are described in Alam et al. 2021.

### 1.5.7 Weak gravitational lensing

Photons follow space-time geodesics. Space-time is distorted in the presence of a source of gravitational potential, which is typically a mass concentration. Therefore, the matter distribution in the Universe bends trajectories of photons travelling from distant sources. This phenomenon receives the name of *gravitational lensing*, since the theory describing light propagation on a gravitational field is analogous to classical optics. Gravitational lensing is a rich cosmological probe since the distortion of photon trajectories depends on both the baryonic and dark matter. From lensing measurements we can learn about the total matter distribution in an expanding Universe.

In the case of an extended source of photons, e.g., a galaxy, light from different locations within the source take slightly different paths towards the observer. If this source is placed in

the background of a mass concentration - the lens - each path will suffer a slightly different bending, which depends on the impact parameter of each photon relative to the center of this lens. Also, due to conservation of surface brightness, the total flux is also increased. The final result is an image that is shifted, distorted and brighter.

Lensing measurements use shifts, distortions and the increase in fluxes to determine properties of the lenses. These measurements are difficult since we do not have access to the original unlensed positions, shapes or fluxes of source galaxies. One remarkable exception is lensing in its strong regime. If the impact parameter of the source relative to the centre of the lens is below some threshold, the Einstein radius, multiple images of the same source are created. Strong lensing allows us to estimate the mass and density profile of a given lens. Moreover, if the source is variable in time, this variability is slightly delayed between each of the multiple images, which allows us to constrain the expansion of the Universe. As mentioned in section 1.5.1, strongly lensed quasars (variable sources) are used to constrain  $H_0$ , under the assumption that we can properly model the lens density profile.

In the weak lensing regime, only statistical measurements can be performed. The main idea is that galaxy shapes can be approximated by ellipses which have an intrinsic distribution of ellipticities and orientations. The matter density field between the observer and source galaxies slightly modifies these shapes and orientations. These changes are observable at the two-point statistics level. Two key ingredients make this measurement challenging: galaxy shapes and redshifts. A weak lensing survey is a photometric survey trying to optimise the sky coverage and image quality. Billions of galaxies have been detected in the state-of-the-art surveys. The shapes of the galaxies are affected by the point-spread function of the images, which is a combination of atmospheric and instrumental dispersion. These effects have to be very well characterised and understood. Galaxy shapes are then converted into shear estimates. A spectroscopic measurement of the redshift for all galaxies is unfeasible with today's technology, so they are estimated using the fluxes measured in a few photometric bands. The precision of the redshift is therefore degraded, typically of the order of  $\sigma_z \sim 0.05(1+z)$ , but since weak lensing is mostly an angular effect, these uncertainties do not play a major role.

With a sample of billions of galaxies with their measured shapes and photometric redshifts, three types of two point functions can be computed. First, the angular two-point correlation function of galaxy positions  $w(\theta)$  as in standard galaxy clustering. Second, the cross correlation between galaxy positions and the shear of galaxies around them. Typically one considers positions from the lenses and the shear of background source galaxies. Also, one considers only the tangential component of the shear relative to each galaxy, where the signal is maximal, in contrast with the radial component. We denote this cross correlation  $\gamma_t(\theta)$ . This second statistics is also referred to as *galaxy-galaxy lensing*. Third, the auto correlation of shear also known as *cosmic shear*. The shear-shear correlations can be projected into two orientations, so they actually give two statistics  $\xi_+(\theta)$  and  $\xi_-(\theta)$ .

The model for the shear of a source galaxy is given by a sum of many layers of matter acting as thin lenses. As in classical optics, the impact of each lens layer depends on the distances between observer-lens and lens-source. The maximum effect is when the lens is at mid-distance between the observer and the source. The model for the two-point functions is therefore an integral of the matter power spectrum, weighted by the distances of the lenses and sources. The angular power spectra are computed and converted into configuration space. Given this dependency on both geometry and the amplitude of the power spectrum, weak lensing measurements are sensitive to the combination defined as  $S_8 = \sigma_8 \sqrt{\Omega_m/0.3}$ .

The most recent measurements have been mainly carried out by three projects. Simi-

larly to the case of the CMB, one large project observed a large area of the sky while two smaller projects observed a smaller patch of the sky to greater depth. The Dark Energy Survey (DES) released recently their cosmological analysis of the Year 3 sample D. Collaboration et al. 2021). They measured 5000 deg<sup>2</sup> of the Southern sky using a 570 Megapixel camera in five photometric bands up to a limiting magnitude of 23. More than 100 million galaxies had their shapes and redshifts measured. This state-of-the-art sample for weak lensing measurements obtained  $S_8 = 0.776 \pm 0.017$  and  $\Omega_m = 0.339^{+0.032}_{-0.031}$  when assuming flat  $\Lambda$ CDM. In a  $w$ CDM model with varying equation of state for dark energy, they obtain  $S_8 = 0.775^{+0.026}_{-0.024}$ ,  $\Omega_m = 0.352^{+0.035}_{-0.041}$  and  $w_0 = -0.98^{+0.32}_{-0.20}$ . The Kilo Degree Survey (KiDS) measured a smaller patch of 1350 deg<sup>2</sup> of the Southern sky in nine optical and infrared bands. Since their area is not as large as DES, they combine their lensing measurements with clustering from the BOSS survey Heymans et al. 2021). The Subaru Hyper Supreme Cam (HSC) is a deep survey to be carried out over 1400 deg<sup>2</sup> in five optical bands up to a limiting magnitude of 26. Results from the first year of data covering 137 deg<sup>2</sup> can be found in Hikage et al. 2019 but their constraints on  $S_8$  are still quite broad compared to those from DES.

Future surveys such as Rubin-LSST and Euclid will significantly improve constraints from weak lensing. Rubin-LSST will observe most of the Southern sky to similar depth as HSC, while Euclid will be the first space program able to produce weak lensing measurements. The great advantage of Rubin-LSST is the large number of optical and near-infrared bands while Euclid does not suffer from atmospheric effects and will have the best galaxy shape measurements. The synergy between these two surveys has been widely explored in the literature (Jain et al. 2015; Rhodes et al. 2017; Capak et al. 2019).

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## Chapter 2

# Observing the Universe with spectroscopy

In the last two decades, spectroscopy became one of the most powerful techniques to survey galaxies across the Universe. Particularly thanks to its capability to obtain precise galaxy redshifts, spectroscopy allows us to build precise maps of the distribution of matter in three dimensions.

This chapter is an overview on how to observe galaxies with spectroscopy and how the data is treated from the target selection all the way to the redshifts. I expose my work on improving the spectroscopic data reduction pipeline for the extended Baryon Oscillation Spectroscopic Survey (eBOSS), for which I was the *Lead Data Scientist* for 3 years.

Naturally, this chapter will focus on the spectroscopic observations with the Sloan Digital Sky Survey (SDSS), but the majority of the concepts introduced here also apply to the Dark Energy Spectroscopic Instrument (DESI).

### 2.1 Selecting the objects to observe

The first step in building a fibre-based spectroscopic survey is to pre-select the objects to be observed. This step, known as *target selection*, is required since one needs to know where to point the optical fibres that take the light from the objects to the spectrographs. Therefore, we cannot simply observe all objects in a given field, we need to choose which targets to observe.

For the target selection, a prior *photometric or imaging* survey is required. In the first years of SDSS, a photometric survey was carried out, covering more than  $14\,555\text{ deg}^2$  of the sky York et al. 2000. The focal plane was equipped with six rows of five charge-coupled devices (CCD), each one covered with one of the SDSS filters: *u*, *g*, *r*, *i* or *z* Gunn et al. 1998; Doi et al. 2010. A technique named drift-scanning was used to continuously observe “stripes” of constant declination during the night. The SDSS was the first of its kind to produce a systematic survey of the Universe in the optical domain, including data releases to the community.

Images were reduced using the SDSS photometric pipeline Lupton et al. 2001; Padmanabhan et al. 2008. Fluxes/magnitudes and their uncertainties were computed for each detected object in five colour bands. Based on their fluxes and angular sizes relative to the point-spread function (PSF), each object received a photometric classification as star or galaxy.

The final list of objects with their respective fluxes and angular positions is the input for targeting algorithms. These algorithms aim to select a given type of object for spectroscopic follow-up, based solely on their fluxes and colours. For galaxy surveys, it is vital to be able

Table 2.1: Surveys and their target selection algorithms

Sample	Redshift range	Reference
SDSS MGS	$0.0 < z < 0.2$	Strauss et al. 2002
BOSS LOWZ galaxies	$0.2 < z < 0.4$	Reid et al. 2016
BOSS CMASS galaxies	$0.4 < z < 0.7$	Reid et al. 2016
BOSS Ly $\alpha$ forest quasars	$2.0 < z < 3.5$	Ross et al. 2012
eBOSS LRGs	$0.6 < z < 1.0$	Prakash et al. 2016
eBOSS ELGs	$0.7 < z < 1.1$	Raichoor et al. 2017
eBOSS quasars as tracers	$0.8 < z < 2.2$	Myers et al. 2015
eBOSS Ly $\alpha$ quasars	$2.0 < z < 3.5$	Myers et al. 2015 Palanque-Delabrouille et al. 2016
DESI BGS	$0.0 < z < 0.6$	Hahn et al. 2022
DESI LRGs	$0.4 < z < 1.0$	Zhou et al. 2022
DESI ELGs	$0.6 < z < 1.6$	Raichoor et al. 2022
DESI quasars	$0.5 < z < 4.0$	Chaussidon et al. 2022

to distinguish between galaxies - the objects of our interest - and stars - which belong to our own galaxy and have a distinct scientific purpose. Additional colour cuts also help selecting a given redshift range for particular types of galaxies.

Since we are interested in the clustering of galaxies, it is essential to obtain a relatively homogeneous angular density of targets so to avoid spurious correlations. Target selection algorithms enforce a requirement of about 15 per cent on the fluctuations of the angular number density of targets. Residual fluctuations have to be corrected before any clustering measurements. I will discuss further about this issue in section 4.1.

Table 2.1 provides a summary of target selection algorithms for several types of galaxy types and redshift ranges.

## 2.2 Pointing fibres to the sky

Once the targets are chosen, we need to define the observing strategy for spectroscopy. This strategy is defined based on several constraints, such as

- the focal plane dimensions, which is a one meter diameter plate holding 1000 optical fibres;
- the field of view of the telescope, which is about  $5 \text{ deg}^2$  for the plates;
- the number of fibres. There are a total of 1000 available fibres of which 80 are used for sky observations and 20 for standard stars;
- the size of the extragalactic footprint, which is roughly  $10\,000 \text{ deg}^2$ ;
- the fibre completeness, i.e., the fraction of targets receiving an optical fibre. The completeness has to be usually above a certain threshold over all the footprint;
- exposure times and total observing time available. Exposure times are dependent on the average signal-to-noise ratio of the observed targets, which need to reach a certain threshold. The total observing time of the program is roughly four to five years.
- visibility window of a given patch of the sky at a given time of the year;

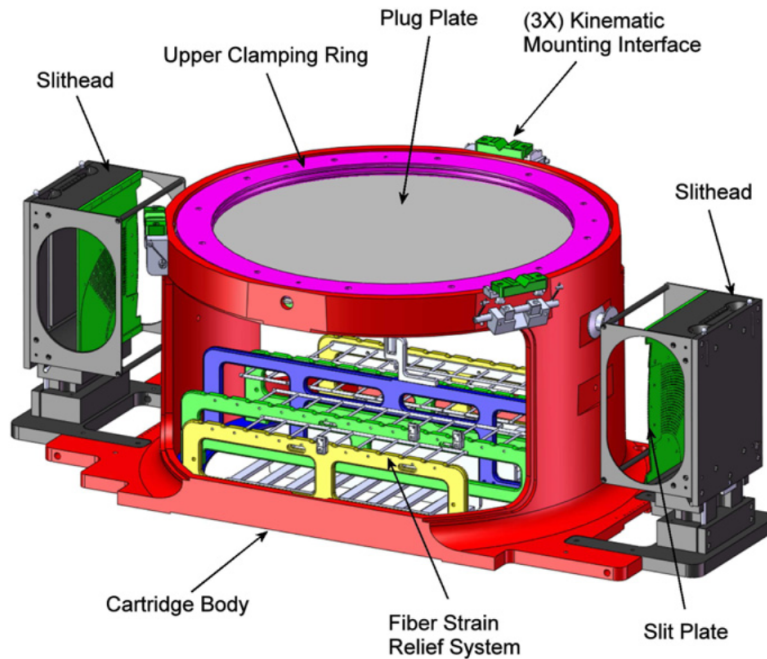


Figure 2.1: Illustration of the SDSS cartridge holding the focal plane plate and the optical fibres. Figure extracted from Smee et al. 2013.

- priority for fibre assignment. Some types of targets have higher priority than others, which affects the fibre completeness of the low-priority samples.

The process of dividing the sky into overlapping projections of the focal plane is called *tiling*. A detailed description of the tiling algorithm can be found in Blanton et al. 2003. Once the tiling and fibre assignment are set, this information is sent to the plate production and drilling of holes that will hold the optical fibres. Focal plane plates are drilled a few months before observing and are unique for a particular patch of the sky and observing time. Not observing the plate at the designed hour angles causes loss of flux due to increased absorption and refraction by the atmosphere.

The drilled plates are sent to the Apache Point Observatory (APO) in New Mexico where the Sloan 2.5-meter Telescope is based. On the mountain, observers attach the plates into cartridges that will fit at the focal plane of the telescope. There are about 15 cartridges, each equipped with 1000 optical fibres. The fibres are plugged by hand, by one or two observers during the afternoon preceding the observation night. Plates are unplugged from their cartridge once a sufficient number of exposures has been taken. A minimum of 3 exposures are taken per plate. Figure 2.1 shows an illustration of a cartridge with the focal plane plate on top.

The light of the objects is transported by the optical fibres through the Sloan spectrographs, shown in Figure 2.2. There are two spectrographs attached at the focal plane of the telescope. Each spectrograph receives the light from 500 fibres and passes it through a beam-splitter, dividing it into a red and blue channels. Each channel has its own grism that spreads the light over wavelength before hitting the CCDs. The blue camera observes roughly from 3500 to 6000Å and the red camera from 6000 to 10500Å. The resolution  $R \equiv \lambda/\Delta\lambda$  increases with wavelength from 1500 to 2000 on the blue camera and from 2000 to 2500 on the red camera.



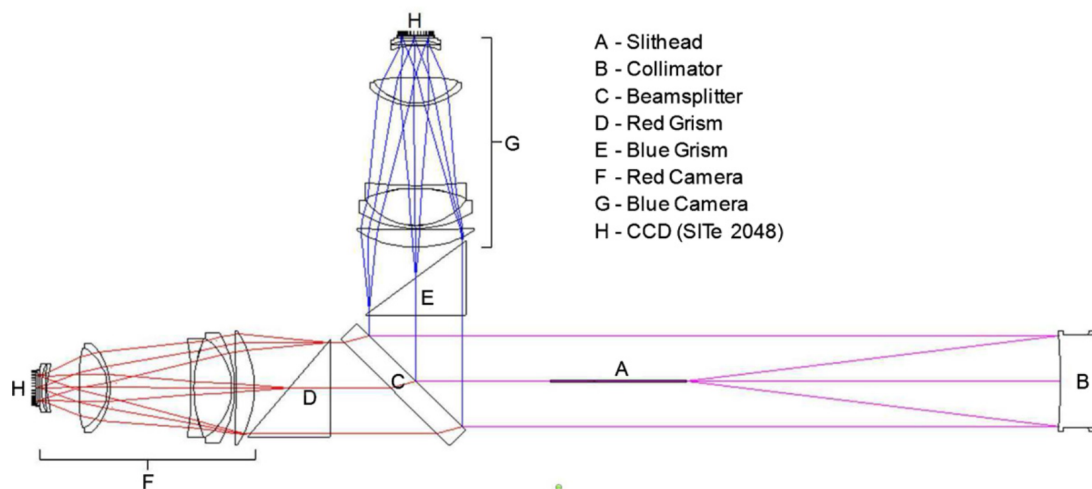


Figure 2.2: Illustration of one SDSS spectrograph. Light propagates from the slithead (also shown in Figure 2.1) towards the blue and red cameras. Figure extracted from Smee et al. 2013.

In addition to the science exposures (observing galaxies), a small set of calibration exposures is also taken, typically at the beginning and at the end of the observing run. Flat exposures are taken with lamps that emit over all wavelengths. The light is passed through the spectrographs, so we refer to these as fibre-flats, as opposed to the exposures taken without the spectrographs, named super-flats. With another type of lamp, which emit narrow lines at some specific wavelengths, arc exposures are obtained, that are used to derive the relation between wavelength and CCD pixel location.

### 2.3 From electrons to spectra

This section describes the data reduction pipeline of spectroscopic observations by the SDSS telescope, for which I contributed as the eBOSS Lead Data Scientist. This automated pipeline transforms the counts stored in CCDs into calibrated spectra, for which redshifts are estimated. Figure 2.3 displays a flowchart of the process. The software, named `idlspec2d`, was written in Interactive Data Language (IDL) and can be found online<sup>1</sup>. The latest version used in Data Release 16 of eBOSS data is `v5_13_0` (Ahumada et al. 2020).

The dispersed light of each object falls onto CCD detectors containing 2048x2048 square  $24\mu\text{m}$  pixels. There are 500 traces per CCD, except when fibres are broken or unplugged by accident. The traces of each spectra are parallel and slightly curved towards the edges. They are separated by about 7 pixels.

The first step of the data reduction is to remove bias and dark, mask cosmic rays and other known bad pixels, convert counts into electrons using estimated gain values, and correct for the super-flat image (flat taken without the spectrograph).

The next step is to extract the total number of counts per wavelength and per object. One of the axis of the CCD is aligned with the wavelengths, but we do not know the wavelength solution at this point. The other axis is aligned with fibre number. The extraction of the fluxes is performed by bundles of 20 fibres. A set of 20 Gaussian profiles plus a third order polynomial

<sup>1</sup>[https://svn.sdss.org/public/repo/ebooss/idlspec2d/tags/v5\\_13\\_0/](https://svn.sdss.org/public/repo/ebooss/idlspec2d/tags/v5_13_0/)

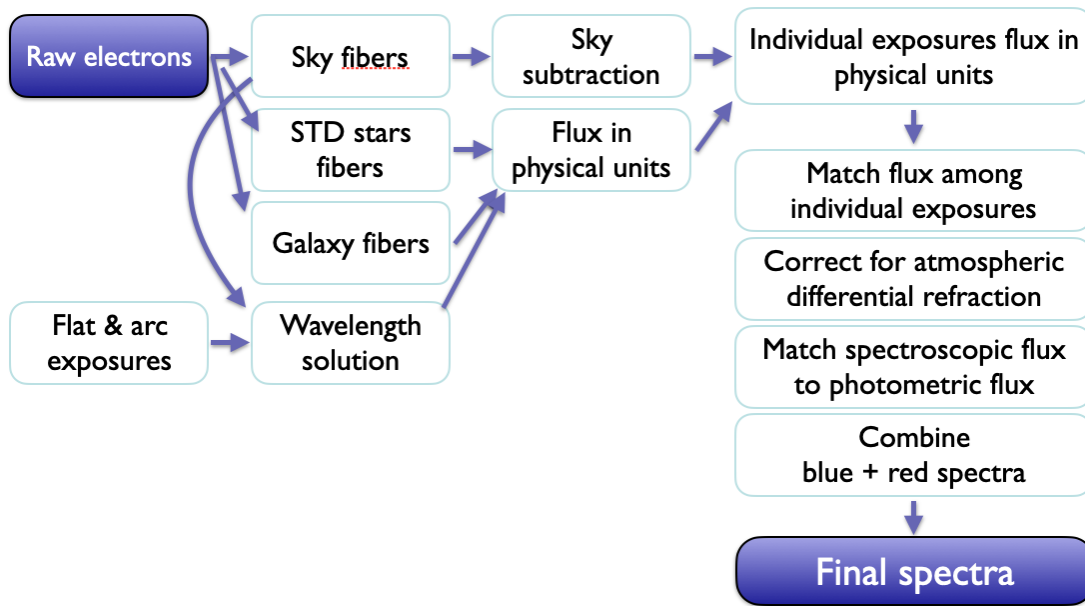


Figure 2.3: Flowchart summarising the several steps of the automated data reduction pipeline from SDSS spectroscopic data.

term are fit simultaneously over the counts. From the best-fit parameters of the Gaussian, one can compute the total flux at each wavelength for each object. This fit is performed regardless of whether the fibres contain flux from sky, stars or galaxies. The extraction step is performed similarly to science, flat and arc exposures.

One important aspect of extraction is: what do we use as weights in the fit? For SDSS-II and III, the extraction used the total estimated variance of each pixel, assumed to be Poisson with mean equal to the number of observed hits in the pixel. However, in SDSS-IV eBOSS we pushed the limits of the instrument by observing fainter objects. In this regime, we started to observe biases due to this weighting scheme in the extraction. The ideal extraction would use the true variance (Horne 1986), not the estimated one, as a weight. Using the estimated one yields a bias in the final fluxes. We modified the extraction algorithm such that it would use a flux-independent weight for the fit, yielding unbiased fluxes. Consequently, this extraction is less optimal, yielding slightly larger flux uncertainties. Biased fluxes affected particularly the analysis of Ly $\alpha$  forests, as described in chapter 3 or in the appendix of Bautista et al. 2017.

The fibre-flat images are used to calculate the traces positions and widths (more precisely than in science images) and to correct for throughput variations across fibres. The arc images are used to calculate the wavelength solution and the dispersion in the wavelength direction based on the line widths. Sky lines in science exposures are eventually used to do small adjustments to the arc-image solution.

The flux in the sky fibres are used to fit a sky model in units of counts. A polynomial dependency over fibre index is used to account to variations over the focal plane. This sky model is then subtracted from all science spectra, including the sky spectra themselves and calibration stars. The sky-subtracted sky fibres are a good metric to evaluate the quality of sky subtraction algorithm.

The next step of the reduction is the *flux calibration*, which converts the observed counts into flux in physical units. The spectra of standard stars are the main ingredient of the flux

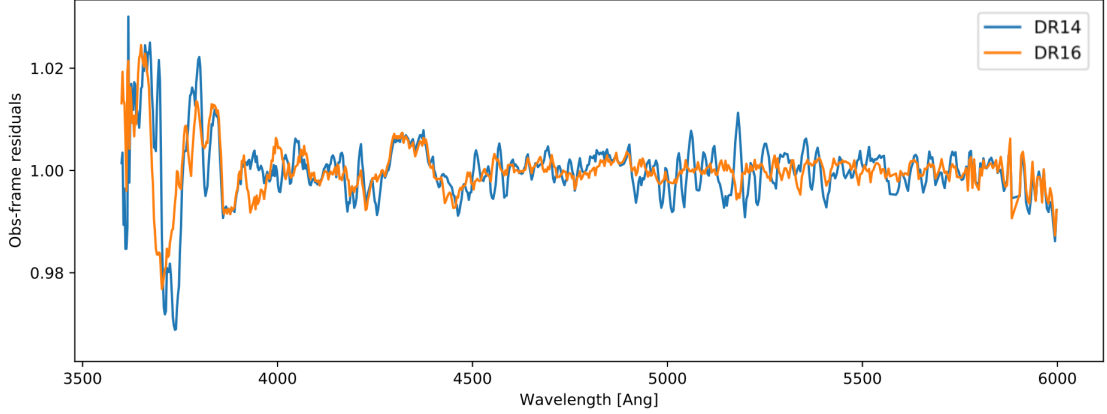


Figure 2.4: Flux calibration residuals computed using an observer-frame stack of quasar spectral regions without emission or absorption lines. Two versions of the automated pipeline are displayed. The DR16 includes an update of the stellar models used to fit spectra of standard stars. We see a significant improvement in these residuals with this change.

calibration. The stars chosen for calibration are of type F, with small variations in temperature and metallicity. Physical models of their spectral emission including absorption features can be obtained via complex stellar synthesis calculations. The goal is to fit absorption lines to the data in order to determine the exact model for each observed star. We start by isolating the absorption lines in the observed spectrum (in units of counts) by fitting a smooth function over its shape, then dividing the whole spectrum by this model. This residual spectrum has only absorption features in a dimensionless scale. The same is done for the physical stellar models. We fit these absorption features to all models to find the best star model for a given spectra. Once the best-fit parameters of the star are found, a calibration vector is constructed by simply taking the ratio of the observed counts to the full model including its smooth component. A set of ten stars are fit independently and a single calibration vector is obtained from them. This final calibration vector is then applied to all other galaxy spectra in order to convert their number counts into flux in physical units. During the last years of eBOSS, I updated the set of physical stellar models from the Kurucz catalogue to those used in DESI (Allende Prieto et al. 2018) which have a larger diversity in stellar parameters and more precise absorption features. This update contributed to a significant reduction of flux calibration residuals computed using stacks of spectral regions of quasars without emission or absorption lines. Figure 2.4 shows this improvement.

Once individual science exposures are converted into physical flux units, we proceed to the co-addition of these exposures into a single set of spectra. For a given object, a B-spline is fitted over all observed spectra, using a new wavelength grid with constant steps in  $\log_{10} \lambda$  of  $10^{-4}$ . The best-fit spline is the final co-added spectrum for this object. The co-addition is made independently for each object. At this stage, we also combine spectra from blue and red cameras into one single spectrum covering the full wavelength range of 3500 to 10500 Å. The last step of this process is the calculation of potential broadband distortions in flux caused by atmospheric differential refraction (ADR). Using information from the plate design, the position of the fibres in the plate and the actual observations (airmass, hour-angle), one can compute a correction vector that is also applied to all spectra.

During the whole reduction, pixels or whole fibres can be masked due to any issues, or when the robustness of flux calculations is compromised. Each pixel has a flux and an uncer-

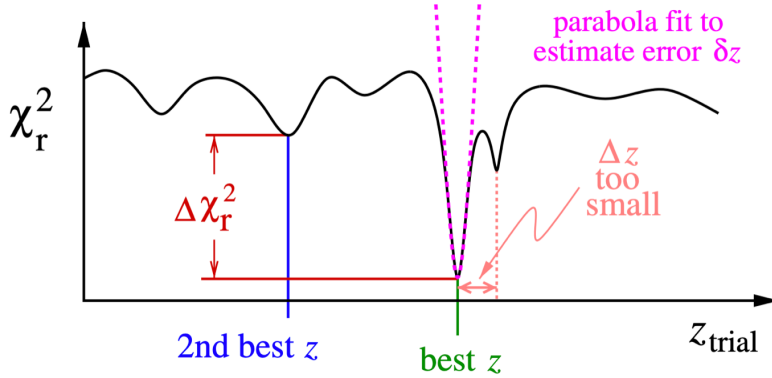


Figure 2.5: Illustration of a  $\chi^2$  profile for a spectral template versus redshift. The best redshift is defined by the location of the global minimum (green). Other minima separated by less than 1000 km/s are not considered to be separate (pink). A parabolic fit to the global minimum (magenta) is used to determine redshift uncertainty. The second-best redshift fit is determined by the location of the second-lowest well-separated minimum (blue). The difference  $\Delta\chi^2$  (red) between best and second-best fits is used to quantify the confidence in the measurement. Figure extracted from Bolton et al. 2012.

tainty estimates. The latter is expressed as an inverse variance, which can be used directly as a weight in analyses.

## 2.4 From spectra to redshifts

The next and last step of the automated data reduction pipeline is the spectral classification and redshift measurement. A set of physical templates of stars, galaxies and quasars is fit to each spectrum by a simple  $\chi^2$  minimization. For each template, we scan over several values of redshifts to obtain the minimum  $\chi^2$ . The redshift ranges probed depend on the type of template: stellar templates are allowed to vary around  $z = 0$ , galaxy templates are fitted over  $0 < z < 2$  and quasar ones over  $0 < z < 7$ . The five best pairs of template and redshift producing the smallest  $\chi^2$  values are stored. Redshift uncertainties are estimated using the  $\chi^2$  profiles around the minima. If the difference between the first and second best-fit  $\chi^2$  values is smaller than a certain threshold (corresponding to roughly  $5\sigma$  for one parameter), a warning flag is set, meaning that the classification is not to be trusted. Figure 2.5 illustrates an example of  $\chi^2$  profile versus redshift, with the potential first and second best-fit redshifts and their uncertainty estimation.

A program of visual inspection of reduced spectra is carried out following the first observations of a new survey. Visual inspection is an important step in order to verify the quality of the data reduction, the automated classification and the redshift measurement. A truth table containing the visually confirmed redshifts and classifications for thousands of spectra is one of the results of the visual inspection. This truth table is used to compute the actual density of tracers that are obtained for a given target selection, as a function of redshift. Spectral features caused by problems in the reduction were then reported to the pipeline team.

During the first months of eBOSS, I coordinated a program of visual inspection where about 15 members analysed about a thousand spectra each (with some overlap for cross-checking). This provided a truth table for the Luminous Red Galaxies from eBOSS. Similar programs were carried out for the eBOSS Emission Line Galaxies and quasars. The inspec-

tions of eBOSS quasars continued throughout the whole program, where a small fraction of the spectra without confident automated classifications were inspected. In DESI, a larger visual inspection program was put in place for the same goals.

This concludes the description of a spectroscopic survey: a machine converting photons to redshifts useful for cosmology.

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## Chapter 3

# The high-redshift Universe and its forests

Until today, the Ly $\alpha$  forest in quasar spectra has been the only tracer of large-scale structures producing measurements of baryon acoustic oscillations (BAO) at redshifts above 2. Given the higher redshifts and the smaller scales probed, the clustering of the forest also yielded competitive constrains on the sum of the mass of the neutrino species, when combining it with measurements of the cosmic microwave background (CMB) anisotropies. Current and future spectroscopic surveys plan to observe denser samples of Ly $\alpha$  forests to improve our understanding of the  $z > 2$  Universe.

In this chapter, I expose my contributions to the study of dark energy through Ly $\alpha$  forest observations. Section 3.1 introduce the main concepts used in this chapter. Section 3.2 focus on the BAO measurements from BOSS and eBOSS surveys for which I provided key contributions. My thesis<sup>1</sup> was dedicated to this topic. Section 3.3 and 3.4 present two projects carried out by my former PhD student Samantha Youles at the University of Portsmouth, UK, who graduated on the 18th of March 2022.

### 3.1 Forests as a tracer of the Universe's structures

The Ly $\alpha$  forest is the name for a series of absorption lines observed in quasar spectra caused by the presence of neutral hydrogen along the line of sight between us and the quasars. Figure 3.1 shows an example of an optical quasar spectrum with its Ly $\alpha$  forest. The forest is seen bluewards of the Ly $\alpha$  emission peak of the quasar at  $\lambda_{\text{obs}} \sim 4900 \text{ \AA}$ , corresponding to a quasar rest-frame wavelength of  $\lambda_{\text{rest}} = 1216 \text{ \AA}$ . The absorption is not limited to the first transition; bluewards of the Ly $\beta$  peak ( $\lambda_{\text{obs}} \sim 4200 \text{ \AA}$  or  $\lambda_{\text{rest}} = 1205 \text{ \AA}$ ) we also observe Ly $\beta$  absorption lines on top of the Ly $\alpha$  ones, and similarly for all Lyman series until the Lyman break at  $912 \text{ \AA}$ .

The advantage of the forests is that it provides a tomographic view of the matter distribution along the line of sight of the quasar. This is because, as the light propagates outwards of the quasar, the Universe expands, causing the whole spectrum to be redshifted. When the redshifted spectrum hits a hydrogen atom, the light being absorbed at the Ly $\alpha$  transition in the *atom* rest-frame is no longer at the Ly $\alpha$  wavelength in the *quasar* rest-frame: it is light from a bluer wavelength. If a given quasar sits at a redshift  $z_q$  and a given intervening atom sits at a redshift  $z_a < z_q$ , the Ly $\alpha$  photon being absorbed by the atom corresponds to a photon at  $\lambda_{\text{rest}} = 1216(1 + z_a)/(1 + z_q) \text{ \AA}$  in the quasar rest-frame. In our frame, the atom absorb

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<sup>1</sup><https://tel.archives-ouvertes.fr/tel-01389967>



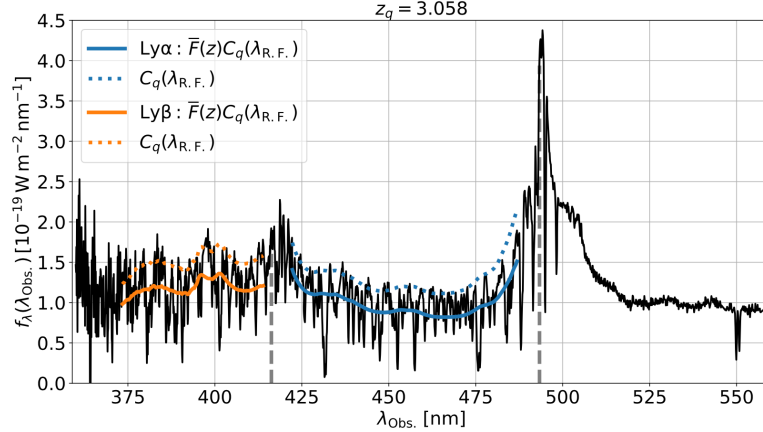


Figure 3.1: Example of a  $\text{Ly}\alpha$  and  $\text{Ly}\beta$  forests from a high signal-to-noise ratio eBOSS quasar spectrum as a function of observed wavelength. The two solid curves give the two independent best-fit models for the product of the mean continuum times the mean absorption. Grey dashed vertical lines show the positions of the  $\text{Ly}\alpha$  and  $\text{Ly}\beta$  broad emission lines of the quasar. Figure extracted from du Mas des Bourboux et al. 2020.

$\text{Ly}\alpha$  photons at the observed wavelength  $\lambda_{\text{obs}} = 1216(1 + z_a)$  Å which does not depend on the quasar redshift. Quasar redshifts are only needed to identify the  $\text{Ly}\alpha$  forest regions in quasar spectra. The result is that a single forest of lines maps the neutral hydrogen density across a large range in redshifts. One single quasar spectrum can yield a high-resolution one-dimensional map of the distribution of neutral hydrogen along its line of sight, across several hundreds of megaparsecs (Mpc).

How much neutral hydrogen is needed to create a forest of  $\text{Ly}\alpha$  absorption lines? Not a lot as it turns out. Given the high cross-section of the  $\text{Ly}\alpha$  resonance, very low densities of neutral hydrogen are sufficient to create a visible absorption line. At redshifts  $2 < z < 4$  where  $\text{Ly}\alpha$  forests are observed in the optical, typical densities found in the intergalactic medium (IGM) are of the order of  $n_{\text{HI}} \sim 10^{-11} \text{ cm}^{-3}$ , enough to absorb a significant fraction of the photons. This HI density corresponds to less than 1 part in  $10^5$  of the total amount of baryons: most of the HI is ionised.

Denser regions completely absorb the light and create saturated lines (zero flux). This is the case of the so-called *high-column density systems* (HCDs), which are often associated with galaxies or proto-galaxies. The extreme cases can be observed with high-resolution spectroscopy, where we can also observe a non-saturated *deuterium* absorption from which we derive constraints on big-bang nucleosynthesis (BBN, see section 1.5.3).

The  $\text{Ly}\alpha$  forest, and therefore the amounts of neutral hydrogen in the IGM, is tightly connected with the process of reionisation of the Universe. As more stars and quasars form, more ultraviolet light is produced, progressively ionising the neutral hydrogen. On top of that, the Universe continuously expands, diluting it. Thus, the average absorption of forests decreases with time (increases with redshift). At redshifts above 6, most of the light in the forest is absorbed, while at redshifts below 1, the  $\text{Ly}\alpha$  absorption is minimal. Forests from  $z_q > 6$  quasars can be used to put constraints on parameters describing the process of reionisation.

The  $\text{Ly}\alpha$  forest and its connection with the large-scale structures have been studied theoretically since works by Gunn and Peterson 1965. Most of recent advances are thanks to numerical simulations. In order to simulate forests, a full hydrodynamic n-body simulations

are required, including all the complex baryonic physics. On large scales, it has been shown that the forest can be considered as linear tracer of the underlying matter field (McDonald et al. 2000; McDonald 2003; McDonald et al. 2005). On smaller scales, the gas physics is more complex to model given the effects of pressure and thermal velocities (Arinyo-i-Prats et al. 2015). Feedback from supernovae explosions, AGN or star forming galaxies also play an important role to model the small-scale clustering (Chabanier et al. 2020). The small scales are interesting due to their potential to constrain neutrino masses for instance.

The first measurements of clustering using the Ly $\alpha$  forests were based solely on the two-point statistics along the same line of sight (Croft et al. 1999; McDonald et al. 2006). Given the high sampling in wavelengths of a forest, the small scales are easily accessible in the radial direction. The smallest scales that can be probed are set by the spectrographs resolution in wavelength. This type of measurement is still performed today and allow us to obtain tight upper limits on the total mass of neutrino species or warm dark matter candidates, when combined with measurements of the cosmic microwave background (Palanque-Delabrouille et al. 2020).

The Baryon Oscillation Spectroscopic Survey (BOSS) from SDSS-III was the first to be able to study the clustering of Ly $\alpha$  forests in three dimensions. Thanks to the density of quasars observed by BOSS, of about 15 deg $^{-2}$ , it was possible for the first time to estimate correlations using absorption from neighbouring lines of sight. Also thanks to the large area of sky covered by BOSS, the first measurement of baryon acoustic oscillations (BAO) using forests was possible Busca et al. 2013; Slosar et al. 2013; Kirkby et al. 2013.

Quasars also trace the matter field. While their density is not homogenous enough across the sky to measure clustering using quasars alone, they can be used to compute cross-correlations with the forests. The cross-correlation between quasars and Ly $\alpha$  forests is interesting because it is mostly independent of the Ly $\alpha$  forest auto-correlation. This is mainly because the quasar sample has a low density (shot-noise dominated) compared to the amplitude of the power spectrum on large scales ( $\bar{n}P_0 \ll 1$ ). The first measurement of BAO in the quasar-forest cross-correlation was also possible with BOSS data (Font-Ribera et al. 2014). The BAO constraints from the auto and cross correlations were combined assuming that these two are independent.

Since the first measurements of BAO using forests, the BOSS and eBOSS collaborations published BAO constraints with increasingly larger samples and improved analysis. They are associated with the official SDSS Data Releases (DR):

- DR9: First measurement of large-scale Ly $\alpha$  correlations without BAO (Slosar et al. 2011);
- DR9: First detection of BAO in the Ly $\alpha$  auto-correlation (Busca et al. 2013; Slosar et al. 2013; Kirkby et al. 2013);
- DR11: First detection of BAO in the quasar-Ly $\alpha$  cross-correlation (Font-Ribera et al. 2014), updated auto-correlation measurement (Delubac et al. 2015);
- DR12: Final BOSS auto-correlation (Bautista et al. 2017) and cross-correlation measurements (du Mas des Bourboux et al. 2017);
- DR14: Updated measurements with eBOSS data (de Sainte Agathe et al. 2019; Blomqvist et al. 2019);
- DR16: Final Ly $\alpha$  BAO measurements of SDSS (du Mas des Bourboux et al. 2020).

My main contributions were given to publications with DR11 and DR12 data.

## 3.2 Baryon acoustic oscillations in the forests

In this section I detail the methodology I used in my past work to measure BAO with a set of Ly $\alpha$  forests, highlighting my personal contributions. I consider the case of SDSS forests, which are observed in the optical domain with low-resolution spectroscopy. Forests have a rather low signal-to-noise ratio on average, so the methods discussed here are adapted to this type of data. Similar methods are employed on the measurement of the line-of-sight (or one-dimensional) power-spectrum, but we focus on BAO here.

The main steps of the BAO analysis with Ly $\alpha$  forests can be summarised in the following steps:

- estimate of the transmission field and their associated weights;
- estimate of the two-point correlation functions, including the cross-correlation with quasars;
- estimate of correction matrices due to distortions of continuum fitting and metals;
- fit of the BAO model over measured correlations.

I will present now in more detail each of these steps.

### 3.2.1 Transmission field

As described in section 3.1, the amount of absorbed flux at a given wavelength (redshift) is related to the density of neutral hydrogen and is therefore a tracer of structures.

The main observable used to compute correlations is the so-called *transmission*  $F$ , defined as

$$F(\hat{n}, \lambda) = \frac{f(\hat{n}, \lambda)}{C(\hat{n}, \lambda)} = \exp[-\tau(\hat{n}, \lambda)] \quad (3.1)$$

where  $f(\hat{n}, \lambda)$  is the observed flux and  $C(\hat{n}, \lambda)$  is the unabsorbed/source flux level, for a quasar line-of-sight at angular position  $\hat{n}$  and observer-frame wavelength  $\lambda$ , which can be linked to the absorber redshift  $z = \lambda/\lambda_\alpha - 1$  if the absorption is due to Ly $\alpha$ . We can also express the transmission as a function of the optical depth  $\tau$  as in the right-hand side of Eq. 3.1.

The challenge is to estimate transmissions from the observed fluxes of a set of quasar spectra, particularly given that our data is low-resolution and relatively noisy. This makes it hard to “see” the unabsorbed flux level  $C$ , also known as the *continuum* level, requiring automated methods. Several past attempts to achieve this goal used either principal-component analysis techniques (Lee et al. 2012), where the templates were built from high-resolution and high signal-to-noise data; or maximum likelihood methods accounting for the non-Gaussian nature of the probability density function of  $F$  (Busca et al. 2013). During my PhD, I particularly tried to merge these last two methods into a single one, without success. Those methods, while more sophisticated and flexible in their modelling of  $C$ , they suffer from the additional noise added to the estimated transmission due to noisier estimates of  $C$ . The simplest method will prevail in the latest analyses, which consists in averaging forests in their rest-frame to obtain a single average shape for  $C(\lambda_{\text{rest}})$  denoted *mean continuum*  $\bar{C}(\lambda_{\text{rest}})$ . This shape is then adjusted onto each individual forest with a linear tilt in  $\log \lambda$ , i.e.,  $C(\hat{n}, \lambda) \equiv \bar{C}(\lambda)[a_0(\hat{n}) + a_1(\hat{n}) \log \lambda]$ , where  $a_0$  and  $a_1$  are fitted coefficients per quasar.

Once the continuum level  $C$  is estimated for each forest, one can compute the transmission  $F$  and their *fluctuations*  $\delta_F$  around the mean, defined as

$$\delta_F(\hat{n}, \lambda) = \frac{F(\hat{n}, \lambda)}{\langle F \rangle} - 1 \quad (3.2)$$

where  $\langle F \rangle$  is the ensemble-averaged transmission of the Universe.

The  $\langle F \rangle$  is actually an evolving function of time (or redshift or observed-frame wavelength) and it is important to take this evolution into account when computing  $\delta_F$ . Several measurements of this quantity exist (Faucher-Giguère et al. 2008; Pâris et al. 2011; Becker et al. 2013; Kamble et al. 2020), though the latest BAO measurements do not use them directly to compute  $\delta_F$ . One could in principle use the forests themselves to estimate  $\langle F \rangle$ , by stacking forests in their observer-frame (by contrast with the mean continuum that is a stack in the quasar rest-frame), though this requires a good estimate of  $C$ . If one express  $\delta_F$  as a function of the observed flux  $f$ , the continuum  $C$  and the average transmission  $\langle F \rangle$ , appear to be degenerate. The latest BAO analyses therefore fit a linear function that models directly the product  $C\langle F \rangle$ .

All methods to estimate the transmission fluctuations also suffer from the fact that they use information from the forests themselves to estimate  $C$  or  $C\langle F \rangle$ . The continuum fitting add extra correlations among pixels in the same forest. As I will discuss in section 3.2.3, these correlations also manifest when considering pixels in different lines-of-sight, distorting the cosmological correlation function. We name this effect the *distortion by continuum fitting* and we attempt to model it when fitting the measured correlation function.

Typical wavelength ranges chosen to define the Ly $\alpha$  forest and extract the transmission field are between 1040 and 1200 Å in the quasar rest-frame. Recent analyses also consider the Ly $\beta$  forest region, between 920 and 1020 Å, where both Ly $\alpha$  and Ly $\beta$  absorption are present. The absorption in this region can be assigned a redshift that depends on the choice of  $\lambda_{\text{rest}}$ , which can be either Ly $\alpha$  = 1216 Å or Ly $\beta$  = 1025 Å. These rest-frame wavelengths are separated enough such that both can be used in clustering measurements without much contamination at separations below 200  $h^{-1}$ Mpc.

Uncertainties of the transmission fluctuations are also estimated from the data themselves. Typically two major components are taken into account when defining uncertainties: instrumental noise and the intrinsic variance of absorbers. The former is essentially the output of the data reduction pipeline, described in chapter 2, corrected with some normalisation factor. The latter is estimated from the variance of  $\delta_F$  observed in the data. The intrinsic variance of  $\delta_F$  is an increasing function of redshift, while the instrumental noise is typically decreasing with observer-frame wavelength. The inverse of the final pixel uncertainty squared is used as a weight in the estimates of the correlation function.

As mentioned previously, high-column density systems (HCDs) leave stronger absorption profiles in the forests and have to be treated differently than standard forest absorption. Several algorithms to automatically find HCDs in noisy data have been developed (e.g. Garnett et al. 2017). Once they are found, the central part of the absorption profile is masked out, while their wings are fitted with a Voigt profile. Only HCDs above a certain column density threshold are efficiently detected, while those below this threshold remain in the forests. The unidentified HCDs increase the large-scale bias of the clustering and alters the amplitude of redshift-space distortions (Font-Ribera et al. 2012; Pérez-Ràfols et al. 2018). Between 15 and 20% of lines of sight contain detectable HCDs.

Broad absorption lines (BALs) are a more complex type of forest contaminant. The BALs are created due to gas flows near the accretion disk of the quasar. These flows create irregular absorption features close to the broad emission lines. Given the complexity of these flows, there are tricky to detect on data but a few algorithms exist. Commonly, forests containing BAL are removed from the analysis, which represents roughly 10 to 15% of the sample.

### 3.2.2 Two-point correlation functions

Once the transmission fluctuations  $\delta_F$  and their weights  $w$  are estimated for all pixels of all forests, we compute their auto-correlation function and their cross-correlation with quasars (or any other point tracer) in three dimensions, as a function of radial and transverse separations.

The first step is to convert redshifts to comoving distances using a fiducial cosmology. With angular positions and distances, we can obtain the separation  $\vec{r} = (r_\perp, r_\parallel)$  between two tracers, where  $r_\perp$  is the component orthogonal to the line-of-sight and  $r_\parallel$  is the component along the line-of-sight.

The correlations are estimated in bins of separation. We consider usually bins of  $4 h^{-1}$  Mpc from 0 to  $200 h^{-1}$  Mpc in both radial and transverse directions. For the cross-correlation with quasars, the radial separation can be negative, meaning that the absorber is closer than the quasar from us. Let  $A$  be the index of a separation bin, the auto-correlation is defined as

$$\xi_A^{\text{auto}} = \frac{\sum_{i,j \text{ if } \vec{r}_{ij} \in A} w_i w_j \delta_{F,i} \delta_{F,j}}{\sum_{i,j \text{ if } \vec{r}_{ij} \in A} w_i w_j} \quad (3.3)$$

where the indexes  $i, j$  denote pixels in the forests. The sum is over all pairs of absorbers for which the separation  $\vec{r}_{ij} = \vec{r}_i - \vec{r}_j$  falls inside separation bin  $A$ .

The cross-correlation of absorbers with quasars is estimated with the following estimator:

$$\xi_A^{\text{cross}} = \frac{\sum_{i,j \text{ if } \vec{r}_{ij} \in A} w_i w_j \delta_{F,i}}{\sum_{i,j \text{ if } \vec{r}_{ij} \in A} w_i w_j}, \quad (3.4)$$

which is very similar to the auto-correlation one in Eq. 3.3. The index  $i$  runs over absorbers while  $j$  runs over quasars, but there is no  $\delta_{F,j}$  term. This estimator is valid under a few assumptions: sparsity of quasars, weak cross-correlation and weak auto-correlation. These assumptions are discussed in detail in the appendix B of Font-Ribera et al. 2012.

The covariance matrix is estimated using sub-samples of the full survey. Under the assumption that each sub-sample  $s$  is independent, the estimator of the covariance between  $\xi_A$  and  $\xi_B$  is

$$C_{AB} = \frac{1}{W_A W_B} \sum_s W_A^s W_B^s [\xi_A^s \xi_B^s - \xi_A \xi_B], \quad (3.5)$$

where  $W_A^s$  is the total weight in bin  $A$  of sub-sample  $s$  and  $W_A \equiv \sum_s W_A^s$ . This estimator has been tested with 100 realisations of synthetic datasets (mocks) and it shows to be robust at the current precision level.

The data vector contains around  $50 \times 50$  bins, making the covariance matrix too large ( $2500 \times 2500$ ) for the usual number of available sub-samples ( $\sim 1000$ ). To avoid the matrix to be singular, we apply a smoothing to it. The smoothing procedure considers that the correlation coefficients, defined as  $\rho_{AB} = C_{AB} / \sqrt{C_{AA} C_{BB}}$ , are only a function of  $\Delta r_\perp \equiv r_{\perp B} - r_{\perp A}$  and  $\Delta r_\parallel \equiv r_{\parallel B} - r_{\parallel A}$ . By averaging all correlation coefficients with the same  $(\Delta r_\perp, \Delta r_\parallel)$ , we obtain a  $50 \times 50$  matrix which is now positive definite. The new covariance matrix is constructed by taking these averaged coefficients and multiplying them by the variances. This method is used to estimate covariance matrices for both auto and cross correlation functions, but also for the cross-covariance between  $\xi_A^{\text{auto}}$  and  $\xi_B^{\text{cross}}$ , used in joint fits (see section 3.2.4).

### 3.2.3 Correction matrices

There are two important effects to be taken into account when modelling the Ly $\alpha$  correlation functions: the distortion by the continuum fitting, discussed in section 3.2.1, and the con-

tamination by metal absorption. Most recent analyses use matrices to convolve a binned cosmological model and obtain the final template to be compared to the measured binned correlations. We describe both corrections below.

The distortion matrix is built under the assumption that the “mistake” we make when fitting the continuum  $C$  is also a linear function of  $\log \lambda$ . We therefore remove completely the components of  $\delta_F(\lambda)$  that are proportional to a constant plus a slope. This procedure removes not only the potential mistake we made, but also the cosmological signal proportional to a constant plus a slope, though it is ok since these components should not contain much BAO information. This removal of a linear component of  $\delta_F$  is referred to as *projection*<sup>2</sup>. Thanks to the linearity of the problem, the model for the correlation function of the projected  $\delta_F$  field can be written as a linear function of the true cosmological correlation function. For the case of a binned correlation function, this relation can be written with the help of a matrix, the *distortion matrix*. The distortion matrix was first introduced in Bautista et al. 2017 for the BAO analysis of DR12 forests, where we showed that the distortion matrix was correctly describing the clustering of mock forests after continuum fitting.

The contamination due to *metal*<sup>3</sup> absorption is mainly due to four transitions of silicon atoms for which their rest-frame wavelength is close to the Ly $\alpha$  transition. When computing the auto-correlation of forests, metal absorption creates confusion because it is interpreted as a distinct Ly $\alpha$  absorber when in fact it is the same absorber but at a different wavelength. For example, the SiIII absorption happens at  $\lambda_{\text{rest}} = 1207 \text{ \AA}$ , close to Ly $\alpha$  for which  $\lambda_{\text{rest}} = 1216 \text{ \AA}$ . In comoving coordinates, assuming a standard fiducial cosmology, this corresponds to  $\Delta r_{\parallel} \approx 21 h^{-1} \text{ Mpc}$  at  $z \sim 2.3$ . At these radial separations and when  $r_{\perp} \sim 0$ , we see a spike in the correlation function. This peak is not due to Ly $\alpha$  absorption separated by  $21 h^{-1} \text{ Mpc}$  but it is the cross-correlation between Ly $\alpha$  and SiIII absorption at near *zero* separation. It is the same region in the Universe absorbing both in Ly $\alpha$  and in SiIII, though the latter has a smaller amplitude. This happens for all metals with rest-frame wavelengths close to Ly $\alpha$ . For instance, one of the SiII transitions creates a spike at  $111 h^{-1} \text{ Mpc}$ , right at the BAO peak location. Since Delubac et al. 2015, I worked on carefully adding metal absorption to mock forests in order to quantify their impact on BAO constraints. In Bautista et al. 2017, we introduced a model for the Ly $\alpha$  -metal cross-correlation to be fitted altogether with the BAO model. Similarly to the distortion modelling, the Ly $\alpha$  -metal correlations can be described as a matrix times the Ly $\alpha$  - Ly $\alpha$  correlation function. This matrix is referred to as the *metal matrix*. The amplitude of Ly $\alpha$  -metal correlations is a free parameter in the BAO fits. We found that marginalising over these amplitudes makes BAO constraints slightly looser, though more conservative and robust.

### 3.2.4 The model

The model used to fit for the BAO scale on the measured correlation functions is based on a linear matter power spectrum with some empirical terms that account for redshift-space distortions, non-linear small-scale clustering, high column density systems and binning effects. The model is constructed in Fourier space and it is Fourier transformed to obtain  $\xi(r_{\perp}, r_{\parallel})$ . The latest version of the model is described in du Mas des Bourboux et al. 2020 and I shortly present it in this section.

<sup>2</sup>This is equivalent to the removal of modes contaminated by photometric systematics in galaxy clustering, as studied in Paviot et al. 2022.

<sup>3</sup>The term *metal* is used in astronomy for elements heavier than helium.

The model can be written, in Fourier space, as

$$P^{\text{model}}(\vec{k}) = b_i b_j (1 + \beta_i \mu_k^2) (1 + \beta_j \mu_k^2) P_{\text{QL}}(\vec{k}) F_{\text{NL}}(\vec{k}) G_{\text{bin}}(\vec{k}), \quad (3.6)$$

where  $\vec{k}$  is the wavevector, with modulus  $k$  and  $\mu_k = k_{\parallel}/k$ ;  $b_i$  and  $\beta_i$  are the linear bias and redshift-space distortions parameters, respectively, and  $i$  is an index referring to absorbers or quasars;  $G_{\text{bin}}$  accounts for the binning of the correlation function,  $F_{\text{NL}}$  is a empirical term that accounts for the non-linear effects on small scales, and  $P_{\text{QL}}$  is the linear matter power spectrum with a empirical anisotropic damping applied to the BAO peak component. The linear matter power spectrum is computed from a Boltzmann solver code, such as CAMB (Lewis et al. 2000).

The non-linear term  $F_{\text{NL}}$  is only included for the quasar- $\text{Ly}\alpha$  cross-correlation and accounts for the effect of Fingers-of-God on small-scales as well as an eventual contribution from uncertainties of redshift estimates, which can be relatively large for quasars ( $\sim 500$  km/s). This term can be modelled as a Gaussian or as a Lorentzian, as a function of a velocity dispersion parameter  $\sigma_v$  (in units of km/s). The impact of redshift errors was studied in detail in Youles et al. 2022 and it is also described in section 3.3.

The effect of high-column density systems, if identified and modelled when fitting the continuum, is negligible. For the HCDs that are not identified, the extended absorption (called “wings”) causes  $k_{\parallel}$ -dependent broadening of the clustering (Font-Ribera and Miralda-Escudé 2012; Rogers et al. 2018). This effect is modelled through effective biases  $b_i^{\text{eff}}$  and RSD parameters  $\beta_i^{\text{eff}}$  that depend on the biases of the  $\text{Ly}\alpha$  forest and those from HCDs, plus a function that depends on the distribution of column densities of the unidentified HCDs.

The metal correlations are accounted as additive terms to the  $\text{Ly}\alpha$  correlations

$$\xi_A^{\text{model}} = \xi_A^{\text{Ly}\alpha \times \text{Ly}\alpha} + \sum_m \xi_A^{\text{Ly}\alpha \times m} + \sum_{m_1, m_2} \xi_A^{m_1 \times m_2}, \quad (3.7)$$

where each  $\xi^{m_1 \times m_2}$  can be written as a metal matrix  $M_{A,B}$  times the  $\xi_B^{\text{Ly}\alpha \times \text{Ly}\alpha}$  as mentioned in section 3.2.3. The metal matrix is defined as

$$M_{A,B} = \frac{1}{W_A} \sum_{(m,n) \in A, (m,n) \in B} w_m w_n, \quad (3.8)$$

where  $(m, n) \in A$  refers to pixel separations computed assuming  $\text{Ly}\alpha$  absorption, while  $(m, n) \in B$  refers to pixels separations computed assuming metal absorption. This formalism was introduced in Blomqvist et al. 2018.

Quasars emit UV photons which ionise their surroundings slightly more than the average intergalactic medium. This is known as the *proximity effect* and it changes the average properties of  $\text{Ly}\alpha$  absorption and therefore its clustering. When considering forests around a given quasar but not to its own forest, we refer to it as the transverse proximity effect. This effect is accounted for in the quasar- $\text{Ly}\alpha$  cross-correlation with the following parametric form:

$$\xi^{TP} = \xi_0^{TP} \left( \frac{1 h^{-1} \text{Mpc}}{r} \right)^2 \exp(-r/r_{UV}) \quad (3.9)$$

where  $r^2 = r_{\parallel}^2 + r_{\perp}^2$ ,  $r_{UV} = 300 h^{-1} \text{Mpc}$  is fixed and  $\xi_0^{TP}$  is a fitted amplitude.

The spectroscopic data reduction pipeline described in section 2.3 also introduces spurious correlations between spectra. For instance, the sky-subtraction and flux calibration processes correlate pixels at the same observed wavelength in two distinct lines-of-sight,

provided they are observed with the same spectrograph (set of 500 fibres). These correlations have been studied using quasar spectral regions deprived of Ly $\alpha$  absorption, such as the MgII spectral region ( $\lambda_{\text{rf}} \in [2600, 2760]$  Å). An empirical model for this effect is included: a Gaussian, function of  $r_{\perp}$ , centred in  $r_{\perp} = 0$  but only when  $r_{\parallel} = 0$ . Due to the effect of distortion caused by continuum fitting, these spurious correlations due to data reduction also propagate to pairs with  $r_{\parallel} \neq 0$ . The amplitude and characteristic length of this effect are two free parameters in the fit of the Ly $\alpha$  auto-correlation function.

The correlation function model accounts for the distortion matrix, as discussed in section 3.2.3. The distorted correlations are written as

$$\xi_A^{\text{dist}} = \sum_{A'} D_{AA'} \xi_{A'}^{\text{model}}. \quad (3.10)$$

It is also possible to add arbitrary smooth functions of separation in order to account for any unexplained correlations and improve the fit to the data. These functions are used to marginalise over any potential information on BAO coming from the full-shape of the correlation functions, which is avoided. This is commonly done in galaxy clustering analysis as I will present in chapter 4.

### 3.2.5 BAO constraints

The last essential ingredient to be added to the model of correlations are the Alcock-Paczynski (AP) parameters. Thanks to the observed baryon acoustic scale, which we know should be isotropic and have a comoving size of  $r_d \sim 105 h^{-1} \text{Mpc}$ , the AP parameters can be well constrained with the data. These AP parameters account for small differences between the assumed fiducial cosmology, used to convert redshifts into distances, and the measured cosmology extracted from the observed BAO peak position.

We commonly use two AP parameters, encoding radial and transverse dilation effects. They can be written as ratios of the observed  $\Delta z_{\text{BAO}}$  and  $\Delta \theta_{\text{BAO}}$  at a given effective redshift  $z_{\text{eff}}$  from equations 1.22 and 1.21 of section 1.5.4, yielding

$$\alpha_{\parallel}(z_{\text{eff}}) = \frac{D_H(z_{\text{eff}}) r_d^{\text{fid}}}{D_H^{\text{fid}}(z_{\text{eff}}) r_d} \quad (3.11)$$

$$\alpha_{\perp}(z_{\text{eff}}) = \frac{D_M(z_{\text{eff}}) r_d^{\text{fid}}}{D_M^{\text{fid}}(z_{\text{eff}}) r_d} \quad (3.12)$$

where  $D_H$  and  $D_M$  are respectively the Hubble and comoving angular diameter distances (Eqs. 1.2 and 1.4).

The AP parameters  $\alpha_{\parallel}$  and  $\alpha_{\perp}$  are used to scale separations when computing the correlation function model as

$$\xi(r'_{\perp}, r'_{\parallel}) = \xi(\alpha_{\perp} r_{\perp}, \alpha_{\parallel} r_{\parallel}), \quad (3.13)$$

and they are let free when fitting correlations. The BAO constraints are defined by the posteriors of  $\alpha_{\perp}$  and  $\alpha_{\parallel}$ , after marginalising over all other free parameters of the model.

Figure 3.2 shows the current BAO constraints from the SDSS DR16 sample of Ly $\alpha$  forests (containing both Ly $\alpha$  and Ly $\beta$  absorption) and quasars (du Mas des Bourboux et al. 2020). The marginalised constraints from the joint fit of auto and cross-correlations are

$$D_H(z = 2.334)/r_d = 8.99^{+0.20}_{-0.19} \quad (3.14)$$

$$D_M(z = 2.334)/r_d = 37.5^{+1.2}_{-1.1} \quad (3.15)$$

$$\rho(D_H/r_d, D_M/r_d) = -0.45 \quad (3.16)$$



which represent a 2.2 per cent BAO measurement in the radial direction and a 3.2 per cent in the transverse direction.

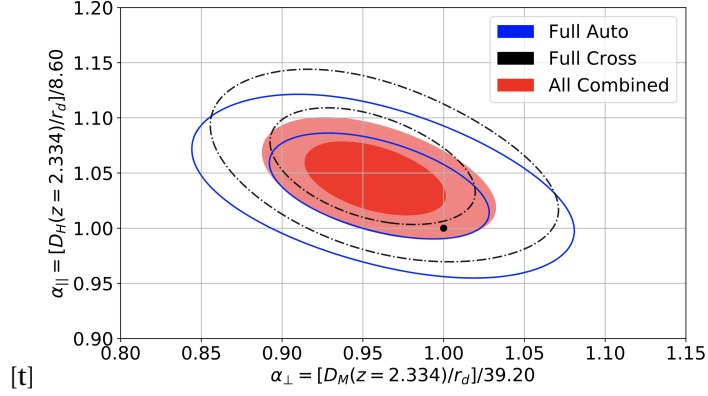


Figure 3.2: BAO constraints from the SDSS DR16 sample of Ly $\alpha$  forests. The contours show the 68 and 95 per cent confidence levels for the Alcock-Paczynski parameters  $\alpha_{\perp}$  and  $\alpha_{\parallel}$  from the fits to the Ly $\alpha$  - Ly $\alpha$  auto-correlation (blue solid), the quasar- Ly $\alpha$  cross-correlation (black, dash-dotted), and the joint fit of auto+cross (red shades). Given the high level of shot-noise of the quasar sample, the covariance between constraints from the auto and cross correlation are found to be negligible. Extracted Figure 12 from du Mas des Bourboux et al. 2020.

### 3.3 Impact of redshift errors

This section concerns work led by Samantha Youles, PhD candidate under my supervision at the University of Portsmouth between 2018 and 2022. The full details are presented in Youles et al. 2022.

Quasar redshifts are a key ingredient in the BAO analysis using forests. Quasar redshifts can be challenging to measure precisely when compared to galaxy redshifts. Quasar spectra are basically composed of thermal continuous emission from their accretion disk around their supermassive black holes in the centre of their host galaxies plus some broad emission lines. Typical quasar spectra from SDSS or DESI do not present detected narrow absorption or emission lines from the host galaxy itself, which would considerably improve the uncertainties on their redshifts. Previous studies using quasars as tracers for clustering (Zarrouk et al. 2018; Lyke et al. 2020) estimated that quasar redshift uncertainties range from 100 km/s at  $z \sim 1$  to 400 km/s at  $z \sim 2$ , while galaxies (LRGs or ELGs) have redshift uncertainties below  $\sim 50$  km/s. Of course, these uncertainties depend on the spectrograph and on which broad emission lines are visible for a given quasar.

In addition to statistical uncertainties in the determination of redshift, Shen et al. 2016 have shown that redshifts derived from broad emission lines of quasars have intrinsic scatter and systematic shifts of the order of hundreds of km/s with respect to the host galaxy redshift. Current redshift fitting algorithms based on templates do not account for these shifts. While Youles et al. 2022 studies the impact of scatter, we left the study of systematic shifts to future work.

The uncertainties in the redshift measurements smear the two-point statistics of the quasars along the line-of-sight. The smearing directly impacts the BAO measurement, particularly reducing the signal-to-noise ratio of estimates of  $\alpha_{\parallel}$  (see section 3.2.4).

Non-linearities induce large peculiar velocities on scales below a few Mpc and also introduce random scatter in the radial direction. These non-linearities are known as *fingers-of-God* (FoG) and they smear the two-point statistics in a similar manner than redshift errors, though with some differences.

In this section, these two types of extra scatter in quasar redshifts are parametrised by

- $\sigma_{v,z}$  as the scatter induced by uncertainties in redshift determination, or redshift errors.
- $\sigma_{v,\text{FoG}}$  as the scatter induced by non-linear motions on small scales, or Fingers-of-God.

Using DESI mock catalogues of the 5-year sample of Ly $\alpha$  forests and quasars, we studied the impact of each of these types of extra scatter on the clustering and on BAO constraints. We discovered a new feature arising from redshift errors that modifies the shape of the two-point statistics in a non-trivial manner, though it does not impact BAO constraints expected with DESI.

Figure 3.3 presents correlation functions, both for the quasar- Ly $\alpha$  cross-correlation and for the Ly $\alpha$  auto-correlation, in averages over  $\mu_r \equiv r_{\parallel}/r$ . There are three sets of mocks

- 1) fiducial ones where  $\sigma_{v,\text{FoG}} = 150$  km/s and  $\sigma_{v,z} = 0$ ,
- 2) mocks with  $\sigma_{v,\text{FoG}} = 500$  km/s and  $\sigma_{v,z} = 0$  and
- 3) mocks with  $\sigma_{v,\text{FoG}} = 150$  km/s and  $\sigma_{v,z} = 500$  km/s.

Each set of mocks contains ten realisations, results are averaged for each set. The colour bands show the scatter amongst realisations and are equivalent to the noise expected for the 5-year sample. We can see that FoG (orange bands) do not have a significant impact except on the cross-correlation on small scales. However, redshift errors do create a new oscillatory feature both in the auto and cross-correlations at separations of 20-50  $h^{-1}$  Mpc but only those along the line-of-sight. Using other mock catalogues with increasing values of  $\sigma_{v,z}$  (not shown here) makes these features stronger. It is the first time this effect is put into evidence.

In Youles et al. 2022, we claim that the oscillatory features appearing when adding redshift errors to quasars is due to two main factors: the smoothing of emission lines in the continua and the auto correlation function of quasars. These two factors *combined* create this new effect, but not individually. The smoothing of the mean continuum is illustrated in Figure 3.4 for increasingly large values of  $\sigma_{v,z}$ . These oscillatory features defined by the difference between true and estimated continua are imprinted on the clustering thanks to the fact that quasars themselves are clustered.

We tested our hypothesis by creating two special sets of mock catalogues, one where quasars are deprived of forest absorption, and one where quasars are not clustered. With the first set of mocks, features are still present with the same amplitudes as the standard mocks. This points to the fact that the Ly $\alpha$  forest and their correlations do not play a role in this effect. For the second set of mocks, the features disappear in the correlation functions, demonstrating that quasar clustering is a key ingredient in explaining the effect.

While this model is satisfactory to explain the features in the cross-correlation, the auto-correlation seems to require an additional ingredient that we could not identify. We leave this investigation for future work.

We performed BAO fits on all sets of mocks for different values of  $\sigma_{v,z}$  and  $\sigma_{v,\text{FoG}}$ . We employ the exact same model described in section 3.2.4, though we did not include metals or high-column density systems. We do not attempt to model the effect of redshift errors when

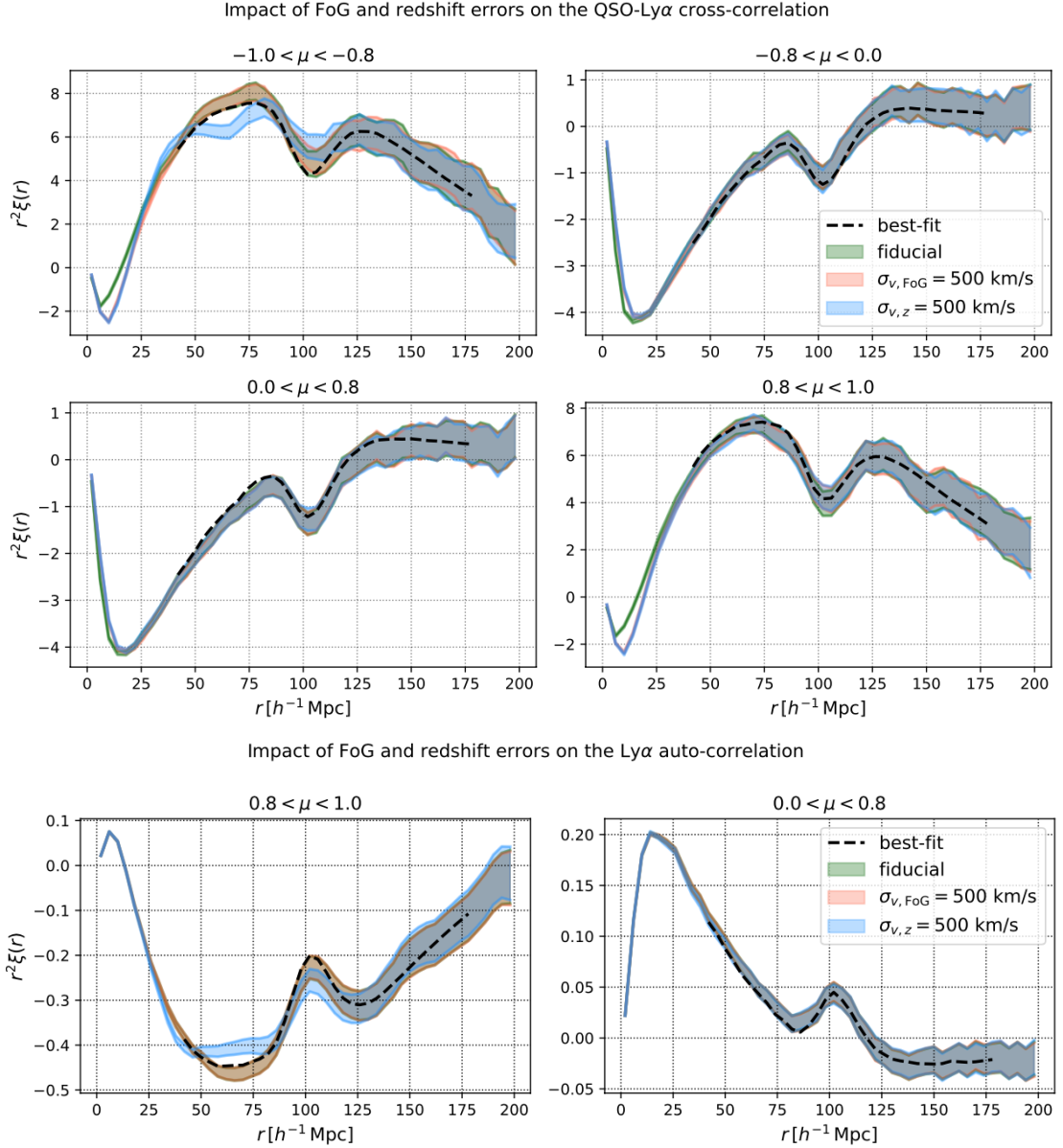


Figure 3.3: Wedges of the correlation functions versus comoving separation. Top four panels show the QSOxLy $\alpha$  cross-correlation (top four panels) while the bottom two panels show the Ly $\alpha$  auto-correlation of forests. These are estimated from DESI mock catalogues reproducing the full 5-year sample. The green bands show the measurement from the standard mocks, with a small value of  $\sigma_{v,\text{FoG}} = 150$  km/s and no redshift errors; the pink bands (indistinguishable from the green in the cross-correlation) correspond to the mocks generated with an extreme value of  $\sigma_{v,\text{FoG}} = 500$  km/s; the blue bands shows the results for mocks with large redshift errors of  $\sigma_{v,z} = 500$  km/s. These measurements are computed from the average of 10 realisations of the complete 5-year DESI survey, and the width of the bands corresponds to the scatter between realisations. Note that  $\mu > 0$  ( $\mu < 0$ ) corresponds to configurations where the Ly $\alpha$  pixels are behind (in front of) the neighbouring quasar.

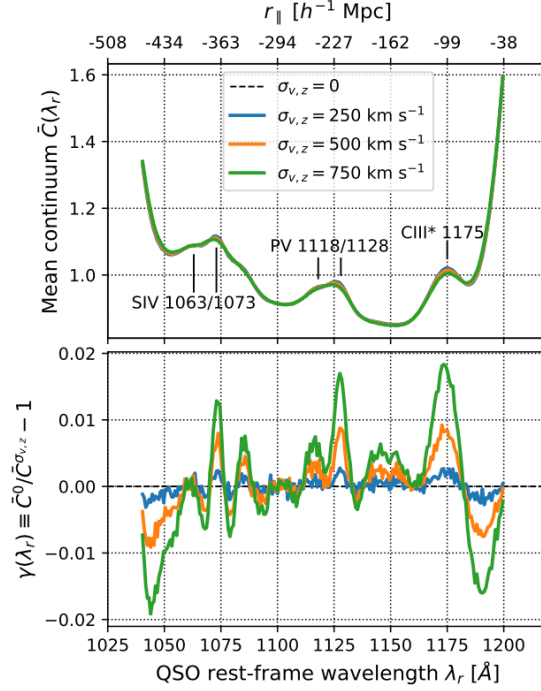


Figure 3.4: Effect of redshift errors on the mean continuum. The top panel shows the mean continua for no-forest mocks without errors, and with increasingly large redshift errors:  $\sigma_{v,z} = 0, 250, 500$  and  $750$  km/s. The positions of some of the strongest emission lines are annotated. The lower panel shows the ratio between the continua with and without redshift errors. The top axis on the top panel shows the radial separations between pixels and the quasar, assuming that the quasar is at redshift  $z_q = 2.3$ .

fitting for BAO, emulating previous analyses. We see an increased value for the  $\chi^2$  of the fits when increasing the value of  $\sigma_{v,z}$ , simply because these features are not accounted for in the model. The inferred BAO parameters  $\alpha_{\parallel}$  and  $\alpha_{\perp}$  do not see significant shifts or trends versus  $\sigma_{v,z}$  other than a slight increase in their uncertainties, which is reassuring.

We believe that, for the final analysis of the 5-year DESI sample, the impact of redshift errors has to be correctly modelled in order to reduce the uncertainties in BAO parameters or to extract cosmological information from the full-shape of the correlation function Cuceu et al. 2021.

### 3.4 Weak-lensing of forests

This section also concerns work led by Samantha Youles, PhD candidate under my supervision at the University of Portsmouth between 2018 and 2022.

Weak-gravitational lensing is a powerful cosmological probe since it is sensitive to the total (baryons and dark) matter density field, as discussed in section 1.5.7. The quasar and Ly $\alpha$  forest samples used in the measurement of BAO lie at  $2.0 < z < 4.0$  and are subject to the effect of lensing from the matter distribution between  $0 < z < 2.0$ . Current BAO analysis such as the one described in section 3.2 do not take any lensing effect into account. In this section we describe how one could potentially extract weak-lensing information from a realistic set of Ly $\alpha$  forests.

Until now, no measurement of weak gravitational lensing with sources at  $z \sim 2$  has been produced, though it has been studied theoretically and with simplified simulations. Croft et al. 2018, Metcalf et al. 2018, and Metcalf et al. 2020 looked into a regular grid of mock quasars with arc minute separations, populated with Gaussian random forests. The close, regular spacing of the forests created a near-continuous field enabling the use of a quadratic estimator adapted from 21cm intensity mapping analyses Pourtsidou and Metcalf 2014; Pourtsidou and Metcalf 2015. Even though these techniques are able to reconstruct the lensing potential, they are not yet readily applicable to real survey data given its non-regular sampling, noise, volumes and densities probed.

Figure 3.5 illustrates the impact of lensing over a pair of Ly $\alpha$  forests caused by an intervening overdensity. The lens creates an image of the pair at a slightly larger angle with respect to the observer, virtually increasing the transverse separation  $r_{\perp}$  between the two forests to a larger value  $r_{\perp}^L$ . Of course, the expected correlation function between two given pixels is unchanged and given by  $\xi(r_{\perp}, r_{\parallel}) = \langle \delta_1 \delta_2 \rangle$  where the brackets indicate an ensemble average. When observing a set of lensed quasars, the measured correlations are wrongly assigned to the separation given by  $(r_{\perp}^L, r_{\parallel}) \neq (r_{\perp}, r_{\parallel})$ . Here we assume that the radial separations changes are negligible,  $r_{\parallel}^L \sim r_{\parallel}$ . If one could measure the correlation function uniquely using pairs of forests that are background of overdense lenses, the results would be a dilated correlation function in the transverse direction. Alternatively, considering only underdense lenses, the effect would be an opposite dilation in the transverse direction.

I describe now how we reconstruct the lensing signal with a new simplified estimator which will be described in Youles et al. (in prep). One should first note that when averaging correlations over all available pairs in a survey, the lensing effect from over and underdense lenses averages to a negligible net impact on the correlation function, though it increases its noise. However, we can attempt to detect a difference between the average correlation function  $\bar{\xi}(r_{\parallel}, r_{\perp})$  and the local correlation function in a given patch of the sky  $\xi(\hat{n}, r_{\parallel}, r_{\perp})$ . We can write a first-order approximate relation between the local difference  $\Delta\xi \equiv \xi(\hat{n}, r_{\parallel}, r_{\perp}) - \bar{\xi}(r_{\parallel}, r_{\perp})$  and the local convergence field  $\kappa(\hat{n})$ . In practice, each pair of forest given us an estimate of  $\xi(\hat{n}, r_{\parallel}, r_{\perp})$ , though very noisy. Our estimates of the convergence  $\hat{\kappa}(\hat{n})$  are assigned into a map by choosing  $\hat{n}$  to be the mid-point between each pair of lines of sight. Convergence maps are constructed using both the auto-correlation function of forests and the quasar-Ly $\alpha$  cross-correlation function, denoted  $\hat{\kappa}_{\alpha\alpha}$  and  $\hat{\kappa}_{q\alpha}$  respectively. Given the fact that auto- and cross-correlations are close to be independent (see section 3.2.2), the convergence maps from auto and cross can be combined into a single map, denoted  $\hat{\kappa}_{q\alpha+\alpha\alpha}$ .

Angular power spectra are computed from the estimated maps as a function of angular multipole number  $\ell$ . Since our convergence maps are quite noisy, computing the auto power spectra of these maps, e.g.  $\langle \hat{\kappa} \hat{\kappa} \rangle_{\ell}$ , is not convenient since these would also be dominated by noise. Instead, we use high signal-to-noise ratio convergence maps to compute cross angular power spectra. In a real setting, one can use a lensing convergence or potential maps derived from the cosmic microwave background or one could construct a convergence map by integrating the density of foreground sources using galaxy survey data.

We tested our methodology using simulations. We used the 5-year DESI mock catalogues of Ly $\alpha$  forests also used in the study of redshift errors (section 3.3). Gaussian lensing convergence maps were created from a fiducial cosmology and the linear matter power spectrum computed with *camb*. We displaced the quasars following these “input” convergence maps and applied our methodology to build convergence maps from the auto- and cross-correlation functions. The analysis was performed 100 times with different realisations of the convergence maps (the forests were kept the same).

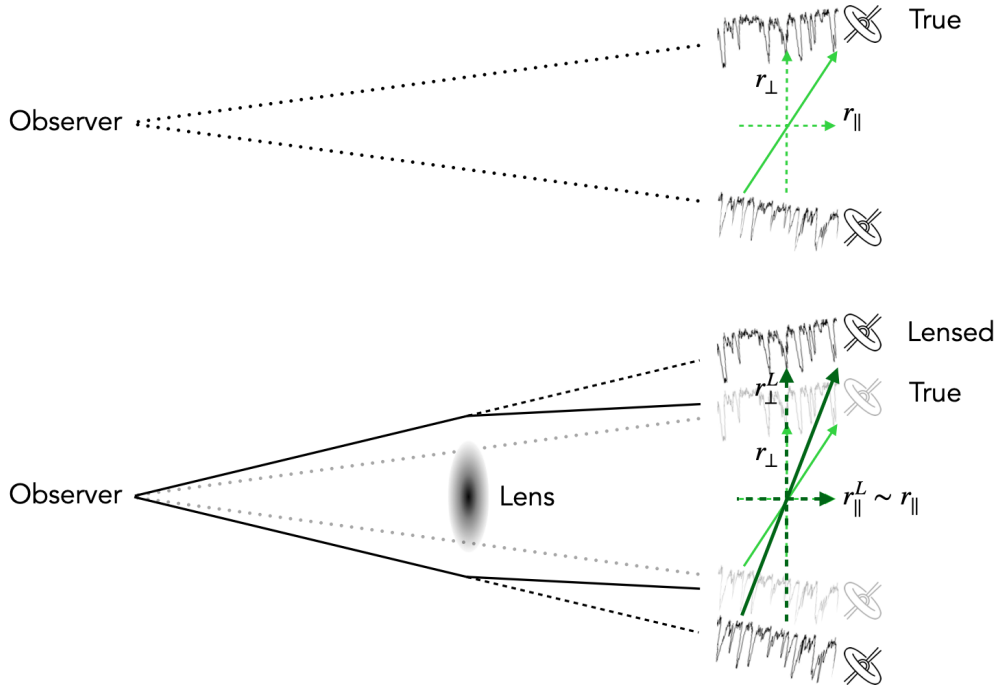


Figure 3.5: Illustration of the effect of gravitational lensing over a pair of Ly $\alpha$  forests. The top figure shows the unlensed case, the arrows define the parallel and transverse separations,  $r_{\parallel}$  and  $r_{\perp}$  respectively, of a pair of Ly $\alpha$  forest pixels (or a pixel-quasar pair). The bottom panel shows the lensed case, where an overdense region acts as a “thin” lens, creating the illusion that the background pairs are actually farther away from each other in the transverse direction. The measured correlations should be the same in both cases, though in the lensed case this correlation is assigned to a wrong transverse separation. The radial effect is relatively small.

Figure 3.6 shows the resulting power spectra compared to the input convergence maps as a function of multipole  $\ell$ . There are two effects to be taken into account when comparing to theory. First, the survey mask changes the amplitude of the power spectra by a factor equal to the fraction of sky covered (to first order). The resulting power is shown as a blue line. Second, the power spectra depend on the choice of maximum transverse separation  $r_{\perp}^{\max}$  used when computing two-point functions used in our convergence estimates. We refer to this non-trivial function of  $\ell$  describing the variation of power with  $r_{\perp}^{\max}$  as “damping”. Both effects combined are shown as a purple line.

There is an excellent agreement, though noisy, between estimated cross-power spectra and their input values. Of course, the “input” map is not available in a real setting so one should consider the noise of this second map (e.g. CMB lensing map) when producing realistic forecasts. We plan to compute realistic estimates on the expected signal-to-noise ratio of these cross power spectra for a DESI like survey of forests.

The results presented in this section are still under development and have not yet been published. We refer the reader to S. Youles thesis for more details on methods and tests.

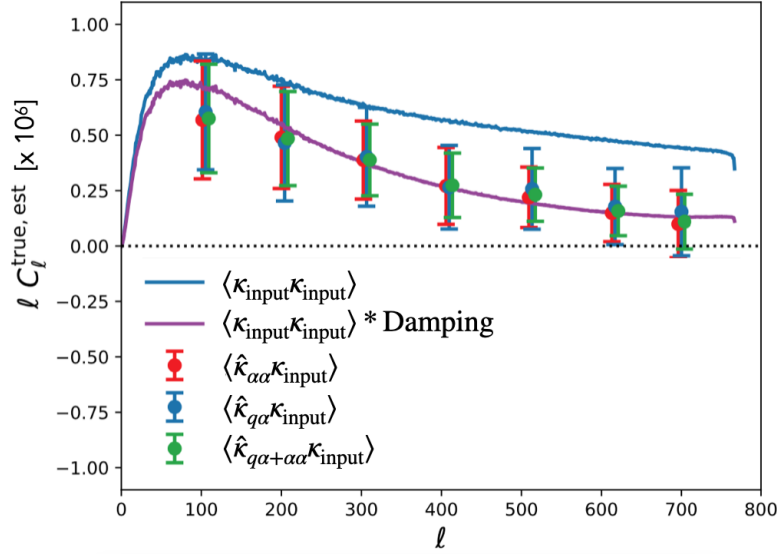


Figure 3.6: Angular power spectra of convergence maps from simulations of the 5-year DESI sample of Ly $\alpha$  forests. The blue line indicates the auto power spectra of the input convergence when accounting for effect of survey mask. The purple line also includes the effect of the choice of maximum transverse separation between lines of sight used when building the convergence maps. The points show the estimated cross power from maps computed from the auto-correlation of forests (red), the cross-correlation between quasars and forests (blue), and the combination of both (green).

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## Chapter 4

# The mid-redshift Universe and its galaxies

At redshifts between 0 and 2, galaxies can be used as tracers of the matter distribution. With the statistics of the galaxy distribution we can measure the characteristic scale of the baryon acoustic oscillations (BAO) as well as extract information about the growth-rate of structures from the anisotropies caused by redshift-space distortions (RSD). As presented in Chapter 1, BAO and RSD are powerful probes of dark energy and theories of gravity.

In this chapter, I overview my contributions for the study of dark energy with galaxy clustering. Section 4.1 describes how to create a galaxy clustering catalogue from the spectroscopic observations, including corrections for all known systematic effects caused by photometry or spectroscopy. Section 4.2 describes reconstruction techniques used to increase the precision of BAO measurements. Section 4.3 presents the mock catalogues that are used to test our methodology and section 4.4 I overview the two-point statistics estimates, either in Fourier space or in configuration space. In section 4.5 I present the BAO measurements I performed with the SDSS sample of luminous red galaxies, while section 4.6 focus on the RSD constraints. Section 4.7 overviews recent work carried out by my PhD students Tyann Dumerchat and Vincenzo Aronica on the joint analysis of galaxy clustering in Fourier and configuration space.

The work described in this chapter is published in the following articles: Bautista et al. 2018; Bautista et al. 2021; Gil-Marín et al. 2020; Ross et al. 2020; Zhao et al. 2021; Dumerchat and Bautista 2022.

### 4.1 From redshifts to clustering catalogues

An essential step in the cosmological analysis of galaxy survey data is to convert a list of galaxy redshifts (see chapter 2) into a catalogue from which we can define overdensities  $\delta_n(\vec{x}) = n(\vec{x})/\bar{n} - 1$ , where  $n(\vec{x})$  is the number density of galaxies in a volume element located at position  $\vec{x}$ . The quantity  $\bar{n}$  is the average galaxy number density over the probed *volume*. Therefore, it is important to 1) define precisely what is the volume observed, which is described by the survey window function; and 2) to ensure that the number density of galaxies is tracing the actual cosmological fluctuations and is not contaminated by other types of fluctuations.

The simplest form of survey window function would be a function of position  $\vec{x}$  that yields 1 if  $\vec{x}$  is inside the survey volume and 0 else. One convenient way to define this function is

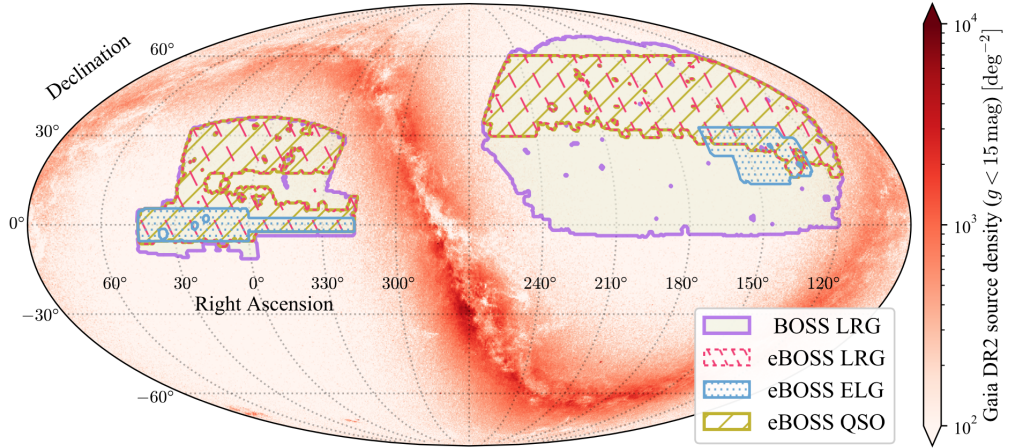


Figure 4.1: The DR16 footprint for each eBOSS tracer: Luminous Red Galaxies (LRG), Emission Line Galaxies (ELG) and quasars (QSO). The BOSS DR12 LRGs area is also shown, as well as the density map of Gaia DR2 sources with  $g < 15$  mag. Figure extracted from Zhao et al. 2021.

to use a large set of randomly and uniformly distributed positions encompassing the survey volume. Each random position would have a weight of either 1 or 0, depending if it is located inside or outside the observed volume. This list of positions and their weights is referred to as the *random catalogue* or simply as the *randoms*. Typically, randoms are sampled at some arbitrary higher number density than the galaxy average number density (typically between 20 or 50 times higher), to reduce the noise in the description of the window function. As we will see next, the random catalogue is slightly more complex: weights are not simply ones or zeros because they account for observational completeness and systematic effects. In SDSS analyses, the random catalogue is can be decomposed into a product of an angular and a radial selection function.

I describe now the procedure to define the angular footprint. The starting footprint is the one used in the process of target selection (section 2.1). Figure 4.1 displays the eBOSS footprint for three tracers (LRGs, ELGs and QSOs) compared to the BOSS footprint and the stellar density map from Gaia. Regions with bad photometry or around bright stars are masked out by removing randoms belonging to these regions (or equivalently, by assigning them a null weight). After tiling, fibre assignment and spectroscopic observations, the footprint can be divided into an unique set of sectors. Each sector corresponds to a connected region observed by a single or several plates. The *fibre completeness* of each sector is defined as the ratio between the number of spectroscopically observed targets to the number of available targets in the sector. Randoms are sub-sampled or de-weighted following the fibre completeness to correct for its effect.

Some targets cannot be observed due to its proximity to another observed target and the physical size of fibres, corresponding to 62 arcseconds in the sky. We refer to these events as *fibre collisions*. Collisions might happen not only to pairs of galaxies, but to any group of targets with linking lengths smaller than 62 arcsec. Depending on the number of plates observing a giving sector, some collisions might be solved. The missing ones impact the measured clustering on small scales if not corrected. While there are several methods to solve collisions (Guo et al. 2012; Bianchi and Percival 2017), we used the simplest up-weighting technique, where the nearest observed galaxy is up-weighted by the number of non-observed targets

within the collision group. This assumes that non-observed targets are also galaxies of the same target type and that they are physically close (angularly and radially) to the observed ones. On scales above a few  $h^{-1}\text{Mpc}$ , this simple correction is a good approximation.

Two additional sets of weights are defined to correct for spurious density fluctuations contaminating cosmological fluctuations in which we are interested. These weights could be applied to randoms but we chose to apply these to the galaxy themselves.

The first set of weights corrects for angular fluctuations caused by correlations between galaxy number density and photometry-related quantities, such as Galactic extinction, stellar density, imaging depth or sky flux. They are commonly referred to as *photometric weights*. Figure 4.2 displays these correlations for the LRG sample both before and after applying correction weights. For the analysis of eBOSS DR16 LRGs, I have implemented<sup>1</sup> a multi-dimension linear regression that assumes these correlations are linear and accounts for the fact that some photometric quantities are correlated between themselves (e.g., stellar density and Galactic extinction). This multi-linear method is based on work described in Prakash et al. 2016. Of course, more advanced methods have been developed since then, based on machine learning algorithms (Rezaie et al. 2020; Chaussidon et al. 2021). It is important to evaluate the performance of these algorithms with mock catalogues where we add artificial contamination and quantify how well we recover the initial raw power spectrum after correcting for them. There is a risk of overfitting and removing some large-scale cosmological modes which can be harmful for clustering analysis, particularly those aiming at measuring primordial non-Gaussianity (Rezaie et al. 2021; Mueller et al. 2021). Vincenzo Aronica is currently working with this issue on DESI mock catalogues, with a particular focus on the low-redshift BGS sample.

The second set of weights corrects for the so-called *redshift failures*, which are spectra with low signal-to-noise ratio (S/N) not yielding a statistically significant redshift measurement. If redshift failures were randomly distributed across the sky, they would not introduce any particular bias to the clustering. However, given the particular configuration and throughput of the SDSS fibres and spectrographs, these failures imprint a pattern in the sky which contaminates clustering. The left panel of Figure 4.3 displays the average failure rate of eBOSS LRGs as a function of position in the focal plane. This pattern is due to the fact that fibres at the side edges of the focal plane transport light to the edges of the CCDs which have a lower-than-average throughput. The right panel of Figure 4.3 shows the dependency of the average redshift failure rate as a function of the signal-to-noise ratio of the observation<sup>2</sup>. When modelling these dependencies of failure rates with S/N (or fibre number or focal plane position), one can correct these spurious fluctuations by assigning a weight inversely proportional to the modelled rate. These *redshift-failure weights* are assigned to the galaxies with confident redshift measurements, to compensate for the eventual non-confident ones at the same location in the focal plane.

Once all galaxies have their correction weights (fibre collision, photometric and redshift-failure) and the randoms are corrected by fibre completeness, it is time to compute the radial selection function. The idea is to assign redshifts to randoms such that they follow the same redshift distribution as the galaxies. The redshift distribution is often quantified by the weighted number density  $\bar{n}(z)$  of galaxies. Figure 4.4 shows the comoving density of eBOSS tracers as a function of redshift. For randoms to match this distribution, we randomly assign

<sup>1</sup>[https://github.com/julianbautista/eboss\\_clustering/blob/master/python/systematic\\_fitter.py](https://github.com/julianbautista/eboss_clustering/blob/master/python/systematic_fitter.py)

<sup>2</sup>The signal-to-noise ratio of an observation was defined internally as a function of the signal-to-noise ratio of all observed spectra and their magnitudes.

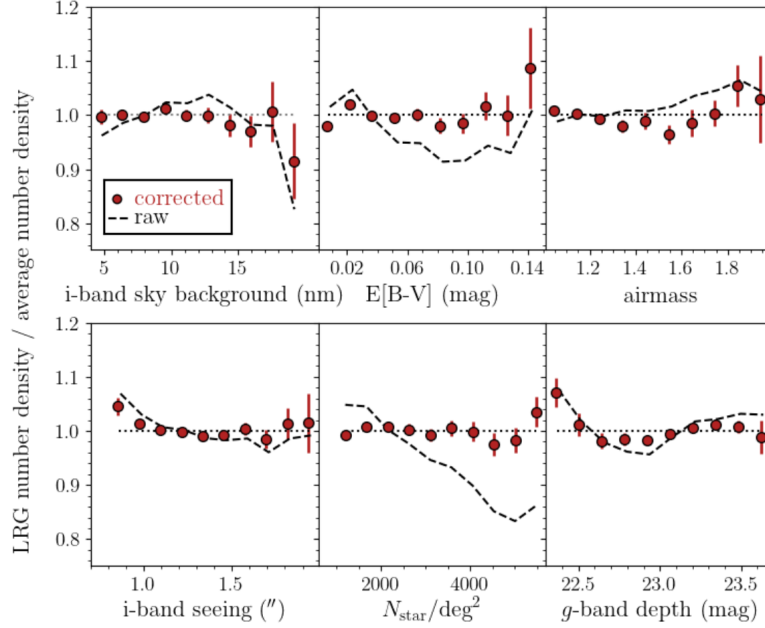


Figure 4.2: Fluctuations in the angular LRG density as a function of various imaging properties and Galactic foregrounds. The dashed curves show these fluctuations before any correction while red points show the result of applying a linear correction for stellar density and Galactic extinction. Figure extracted from Ross et al. 2020.

galaxy redshifts to each random, with repetition. Another possibility would be to fit some smooth function over the observed  $\bar{n}(z)$  and draw random redshifts from the resulting model. All these methods introduce the so called *radial integral constraint*, i.e., the fact that the observed  $\bar{n}(z)$  is not the true ensemble averaged function of the Universe, but derived from the volume-limited realisation that we observe. The impact on clustering of the radial integral constraint is inversely proportional to the area of the footprint (de Mattia and Ruhlmann-Kleider 2019) and it is a significant systematic effect for surveys such as the eBOSS ELG (see Figure 4.1 and Tamone et al. 2020; de Mattia et al. 2021).

The final step of the catalogue production consists in adding an extra set of weights to galaxies that optimise the signal-to-noise ratio of clustering measurements when the redshift distribution is not uniform. These weights were first introduced by Feldman et al. 1994 and have been applied to all clustering analysis ever since. They are commonly referred to as *FKP weights* and can be written for a galaxy  $i$  as

$$w_{i, \text{FKP}} = \frac{1}{1 + \bar{n}(z_i)P_0} \quad (4.1)$$

where  $P_0$  is a rough value for the power spectrum of the tracers at some particular scale of choice, usually BAO scales for large-scale clustering measurements. For eBOSS analyses, we chose a scale of  $k \sim 0.1 h \text{Mpc}^{-1}$  for which  $P_0 = 10000 h^{-3} \text{Mpc}^3$  for LRGs,  $P_0 = 4000 h^{-3} \text{Mpc}^3$  for ELGs and  $P_0 = 6000 h^{-3} \text{Mpc}^3$  for QSOs.

In summary, the random catalogue describes the angular and radial selection functions of the galaxy survey. Each random has a weight  $w_r = w_{\text{comp}} w_{\text{FKP}}$ , where  $w_{\text{comp}}$  accounts the spectroscopic completeness of targets and  $w_{\text{FKP}}$  are optimal weights for clustering. Each galaxy has a weight given by  $w_g = w_{\text{col}} w_{\text{photo}} w_{\text{fail}} w_{\text{FKP}}$ , where  $w_{\text{col}}$  accounts for fibre collisions,  $w_{\text{photo}}$  corrects fluctuations caused by photometry and  $w_{\text{fail}}$  corrects redshift fail-

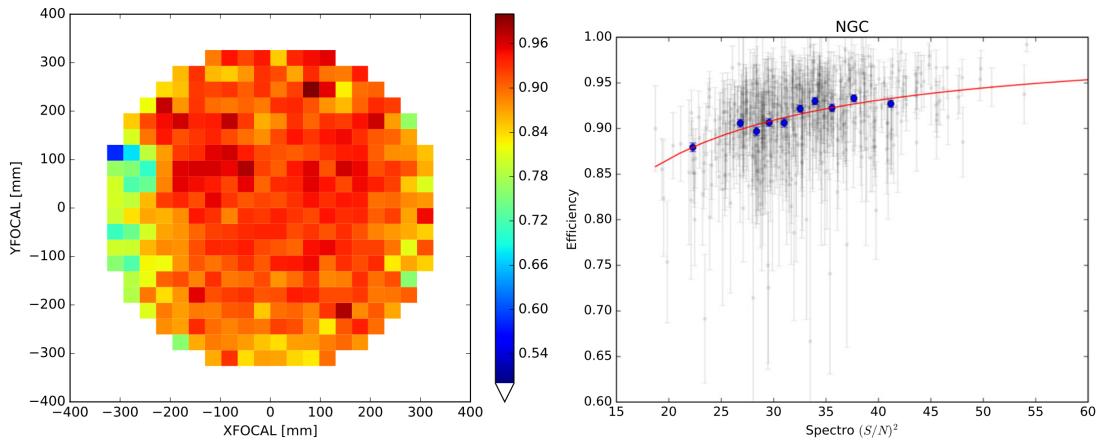


Figure 4.3: Average redshift efficiency as a function of physical position of the optical fiber in the focal plane (left panel) and as a function of the observation signal-to-noise squared (right panel). The left panel shows how redshift-failures imprint a pattern on the sky and therefore must be corrected. The right panel shows the underlying reason for the failures. The model (red line) is used to compute correction weights. Figures extracted from Bautista et al. 2018.

ures. These weights are employed when computing correlation functions or power spectra (see 4.4).

## 4.2 Reconstruction of linear density field

Bulk motions of galaxies on scales of tens of Mpc impact their two-point statistics, acting as a smoothing of the clustering (Eisenstein et al. 2007b). In particular, these motions smooth/broaden the BAO peak, reducing the precision in our measurement of its position. In the past decades, several methods have been proposed to reconstruct the linear density field and remove these bulk motions from a galaxy survey, therefore sharpening the BAO peak. These methods are referred to as *reconstruction* methods. Reconstruction is now an essential part of standard BAO measurements from galaxy surveys, since they significantly improve their precision.

The simplest form of reconstruction is based on a theoretical relationship between the galaxy density field on large scales (our observable) and the displacement field connecting the observed field (Eulerian frame) to its past version (Lagrangian frame). Lagrangian perturbation theory (LPT) can predict these displacements (see Bernardeau et al. 2002 for a review). To first order, the relation between displacements and density is linear (Zel'dovich 1970). From the observed density field in our survey, we can compute the corresponding linear displacements and move galaxies back. Since we want to “move” the density field, we also move randoms by the same displacements. This procedure successfully removes most of the broadening of the BAO peak (Nusser and Dekel 1992; Eisenstein et al. 2007a). The first order LPT reconstruction is often called *Zeldovich reconstruction*. Extending to second order LPT does not improve BAO reconstruction significantly (Seo et al. 2010) but it helps on recovering the velocity field (Kitaura et al. 2012). Zeldovich reconstruction has been used in all BAO measurements from SDSS (BOSS and eBOSS). I implemented a python version<sup>3</sup> of the Zeldovich reconstruction that uses Fast Fourier Transforms while accounting for the large angle effects (Burden et al. 2014; Burden et al. 2015). Alternatively, displacements can

<sup>3</sup>[https://github.com/julianbautista/eboss\\_clustering/](https://github.com/julianbautista/eboss_clustering/)



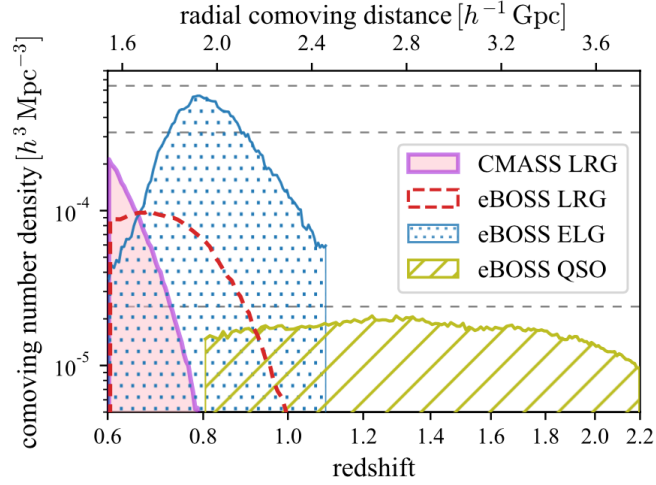


Figure 4.4: The weighted comoving number densities  $\bar{n}(z)$  of eBOSS DR16 tracers and BOSS DR12 CMASS LRGs, with all the photometric and spectroscopic systematic weights included. The comoving distances and volumes are computed assuming a flat- $\Lambda$ CDM cosmology with  $\Omega_m = 0.31$ . Figure extracted from Zhao et al. 2021.

be computed in configuration space using finite difference approximations (Padmanabhan et al. 2012) though these are slower and less precise than iterative FFTs.

Since we do not observe the actual density of matter, a bias relation between galaxy number density and matter density has to be assumed in reconstruction. Usually we assume a linear bias relation, where the bias value is estimated from the clustering of the standard galaxy catalogue. This is a reasonable approximation since reconstruction uses fluctuations on scales larger than  $\sim 15 h^{-1}\text{Mpc}$ . It is also possible to remove linear redshifts-space distortions, by assuming a value for the growth-rate of structures  $f$  (see section 1.5.6) used to convert the displacements into velocities. The radial components of the estimated velocities are then removed from each galaxy. Therefore, the clustering of reconstructed catalogues depend on the fiducial values of bias and  $f$ , but less on the distance-redshift relation (Carter et al. 2020).

There are other backward reconstruction techniques such as those based on the least action principle (LAP). First proposed by Peebles 1989, it was extended to cosmological applications by Nusser and Branchini 2000; Sarpa et al. 2019. Sarpa et al. 2021 applied it for the BAO measurement of the BOSS DR12 galaxy sample.

Elena Sarpa, postdoctoral researcher in my group at CPPM, is currently leading reconstruction efforts in the context of the Euclid Collaboration. Also, she is working on the impact of neglecting massive neutrinos in Zeldovich reconstruction, both at the level of BAO constraints and at the theoretical level. Results from this study could help improving reconstruction in the presence of massive neutrinos and improve mock production methods based on Zeldovich LPT.

Reconstruction is essential to improve the precision on BAO measurements. However, due to imperfections in recovering the true initial conditions, reconstruction has not yet been used in attempts to derive cosmological constraints from the full-shape of the reconstructed two-point statistics. Therefore, analysis of redshift-space distortions are based on the non reconstructed clustering (see section 4.6).

### 4.3 Mock catalogues

An important limitation in cosmology is that we only have a single realisation of our Universe. Therefore simulations are an essential tool to overcome this limitation.

We can recognise two types of simulations mainly used in cosmology. First, n-body simulations solve numerically for the formation of structures from a set of initial conditions, but these are computationally expensive. They focus on the fully non-linear aspects and require a higher resolution setting, preventing them to simulate very large volumes or producing numerous realisations. N-body simulations are still the baseline for tests of perturbation theory models. See Angulo and Hahn 2022 for a review on cosmological n-body simulations. The second type of simulations are the so-called *mock catalogues* or simply *mocks*. They are lower resolution and do not fully describe the non-linear clustering but they are faster to produce both in large volumes and in large number of realisations. They are key for the understanding of biases in best-fitted quantities and on their uncertainties in clustering analyses. Given the large number of realisations, mocks have been the most reliable method to estimate covariance matrices of clustering measurements. Also, mocks are used to study the impact of observational systematic effects on our measurements. All previous cosmological measurements from galaxy surveys have associated sets of n-body simulations and mock catalogues. I will describe the example I was involved in, the case of EZMOCKS produced for the eBOSS DR16 clustering measurements.

A set of 1000 EZMOCK realisations of the eBOSS survey were produced for three tracers: LRGs, ELGs and QSOs. They are fully described in Zhao et al. 2021. These mocks employ an effective Zeldovich approximation to evolve the density field in longer time steps (Chuang et al. 2015). By construction, these mocks match the large-scale clustering of the target data sample, including 2 and 3-point statistics. They have some extra free parameters to simulate a biased galaxy sample and non-linear motions on small-scales. The eBOSS EZMOCKS also include light-cone effects produced by interpolating snapshots at different redshifts.

The last step in the production of EZMOCKS was the addition of systematic effects. I was responsible of implementing all known observational effects from the data onto the mocks. A simple script<sup>4</sup> gathers the information from the real catalogues and simulates them. This process adds fibre collisions, spurious fluctuations caused by photometry and spectroscopy to mocks. These contaminated mocks are then corrected using the same algorithms as for the real data, as described in section 4.1. The clustering before and after adding those effects is extensively compared to real data (Zhao et al. 2021) and the agreement is found to be good on the scales of interest for BAO and RSD measurements.

### 4.4 From catalogues to clustering estimates

Once the catalogue of galaxies and the associated random catalogue are ready, any statistical quantity can be computed, such as  $n$ -point statistics. As discussed in section 1.4, these statistics can be computed in either configuration space or Fourier space. This section describes how to obtain 2-point function estimates, in both spaces, which are the statistics I worked with.

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<sup>4</sup>[https://github.com/julianbautista/eboss\\_clustering/blob/master/bin/ezmocks\\_add\\_systematics.py](https://github.com/julianbautista/eboss_clustering/blob/master/bin/ezmocks_add_systematics.py)

#### 4.4.1 Correlation function

The estimate of the correlation function  $\xi(\vec{r})$  is based on counting pairs of galaxies in bins of separation  $\vec{r}$  and comparing it to same numbers obtained from the random catalogue. For the estimator we use, there is no need to calculate galaxy number densities  $n(\vec{x})$  or overdensities  $\delta_n(\vec{x})$  in some mesh. To optimise the variance of the estimator while keeping it unbiased, Landy and Szalay 1993 proposed to include the cross-pairs between galaxies and randoms in the estimator, yielding:

$$\xi(\vec{r}_i) = \frac{DD(\vec{r}_i) - 2DR(\vec{r}_i)}{RR(\vec{r}_i)} + 1, \quad (4.2)$$

where  $DD(\vec{r}_i)$  (or  $RR(\vec{r}_i)$ ) are the normalised paircounts between galaxies (or randoms) at bin  $i$  corresponding to separation  $\vec{r}_i$ ,  $DR(\vec{r}_i)$  is the normalised cross-paircounts between galaxies and randoms. The normalisation of the paircounts is given by the total number of unique pairs in the catalogue, i.e.,  $N(N-1)/2$  if the catalogue contains  $N$  galaxies or  $N_g N_r$  for the cross pairs. When using weights as described in section 4.1, nothing changes except that each pair of galaxies (or randoms) are weighted as  $w_{ij} = w_i w_j$ , where  $w_i, w_j$  are the weights of each element of the pair.

More recently, new ways to account for fibre collisions or incompleteness have been proposed (Bianchi and Percival 2017; Percival and Bianchi 2017) where each pair of galaxies would have a weight  $w_{ij} \neq w_i w_j$ , i.e., not a product of individual weights. This makes the calculation slightly slower since these weights have to be defined for each pair. These new estimators were first used with the eBOSS tracers (Mohammad et al. 2020) and are currently being employed within DESI.

The estimator in Eq. 4.2 changes when dealing with catalogues resulting from reconstruction algorithms, where both galaxies and randoms are shifted in order to reduce the impact of bulk motions (see section 4.2). The Landy-Szalay estimator is now written as

$$\xi^{\text{recon}}(\vec{r}_i) = \frac{DD(\vec{r}_i) - 2DS(\vec{r}_i) + SS(\vec{r}_i)}{RR(\vec{r}_i)}, \quad (4.3)$$

where  $S$  represents the shifted random catalogue and the  $R$  is the usual random catalogue as before. See Padmanabhan et al. 2009; Padmanabhan et al. 2012 for a detailed derivation of this estimator.

Most analyses of the two-point correlation function do not use the full  $\xi(\vec{r}_i)$ , where  $\vec{r}_i$  are bins in 2D separations such as bins in  $(r_{\parallel}, r_{\perp})$  or  $(r, \mu_r)$ . Most analyses further compress the estimated correlation function into a new basis with fewer degrees of freedom. Some analyses uses the correlation function *wedges*  $\xi_{\mu_{r1}, \mu_r}(r)$ , which are averages of  $\xi$  over  $\mu_{r1} < \mu < \mu_{r2}$ , as already discussed in section 3.2.2. Other analyses, such as the one described in this chapter, decompose  $\xi$  into a basis of Legendre polynomials  $L_{\ell}(\mu_r)$  such that  $\xi(r, \mu_r) = \sum_{\ell} \xi_{\ell}(r) L_{\ell}(\mu_r)$ , where  $\xi_{\ell}(r)$  are the correlation function *multipoles*. Usually two or three wedges or multipoles are considered when comparing to clustering models.

One of the most recent codes to compute correlation function accounting for all features described above is PYCORR<sup>5</sup> which is being currently developed for DESI.

#### 4.4.2 Power spectrum

The power spectrum is the analogous of the correlation function in Fourier space. It requires a Fourier Transform of the galaxy overdensity field, which introduces some effects that have

<sup>5</sup><https://github.com/cosmodesi/pycorr>

to be correctly taken into account when fitting for clustering models. The power spectrum estimation is based on work by Feldman et al. 1994; Bianchi et al. 2015; Hand et al. 2017 and it is shortly described here.

The first step is to assign galaxies and randoms into a three-dimensional mesh, defining an “overdensity” function  $F(\vec{x}_i)$  as

$$F(\vec{x}_i) = n(\vec{x}_i) - \alpha n_{\text{rand}}(\vec{x}_i), \quad (4.4)$$

where  $n$  and  $n_{\text{rand}}$  are the *weighted* number density of galaxies and randoms in a voxel centred on  $\vec{x}_i$ ,  $\alpha$  is the ratio of the total weighted number of galaxies to the total weighted number of randoms. The field  $F(\vec{x})$  is also known as the FKP field.

The field  $F(\vec{x})$  can be Fourier transformed numerically using Fast Fourier Transform (FFT) algorithms. The power spectrum would be defined as the absolute value squared of this field. Power spectrum *multipoles* are commonly used, as it was the case in configuration space. They can be expressed as:

$$P_\ell(k) = \frac{1}{I_{22}} \left[ (2\ell + 1) \left\langle F_0(\vec{k}) F_\ell(-\vec{k}) \right\rangle_{V_k} - (1 + \alpha) I_{12} \delta_{0\ell} \right] \quad (4.5)$$

where the brackets  $\langle \cdot \rangle$  indicate average over spherical shells in k-space with radius  $|\vec{k}|$ ,  $F_\ell(\vec{k})$  are the overdensity multipoles in k-space given by

$$F_\ell(\vec{k}) = \int d^3x F(\vec{x}) e^{i\vec{k}\cdot\vec{x}} L_\ell(\hat{k} \cdot \hat{x}), \quad (4.6)$$

and the normalisation factors  $I_{ab}$  are defined as

$$I_{ab} \equiv \int d^3x [n(\vec{x})]^a [w_{\text{FKP}}(\vec{x})]^b. \quad (4.7)$$

The first term in Eq. 4.5 is the usual term defining the power spectrum, i.e., the square of the absolute value of Fourier fluctuations, while the second term is the so called *shot-noise* and only affects the power spectrum monopole.

Since we use Fourier transforms which rely on periodic boundary conditions, the mesh used to compute  $F(\vec{x})$  has to be large enough to encompass the survey volume with some padding around it. Also, the estimated power spectra will be convolved by the window/selection function of the survey. This window has to be correctly computed using the randoms and it is a key ingredient to be accounted for when modelling the clustering.

Since we use numerical FFTs, the resolution of the mesh defines the scales being probed. If the mesh has a total comoving size of  $L$  on a side and has  $N$  voxels on a side, the fundamental mode is given by  $k_F = \pi/L$  while the highest (or Nyquist) frequency is  $k_{\text{Ny}} = \pi N/L$ . Typically, meshes with 512 or 1024 voxels on a side are sufficient to cover the scales of interest for BAO or RSD analyses in current survey volumes,  $0.01 < k < 0.30(\text{Mpc}/h)^{-1}$ .

The estimated power spectra are also affected by the discreteness nature of the mesh on scales close the Nyquist frequency. The assignment of galaxies to the mesh can be optimised to reduce those effects. Also, interlacing techniques can also further reduce mesh effects. These effects and proposed solutions are described in detail in Sefusatti et al. 2016 and are implemented in the package NBODYKIT<sup>6</sup> or PYPOWER<sup>7</sup>.

<sup>6</sup><https://nbodykit.readthedocs.io/>

<sup>7</sup><https://github.com/cosmodesi/pypower>

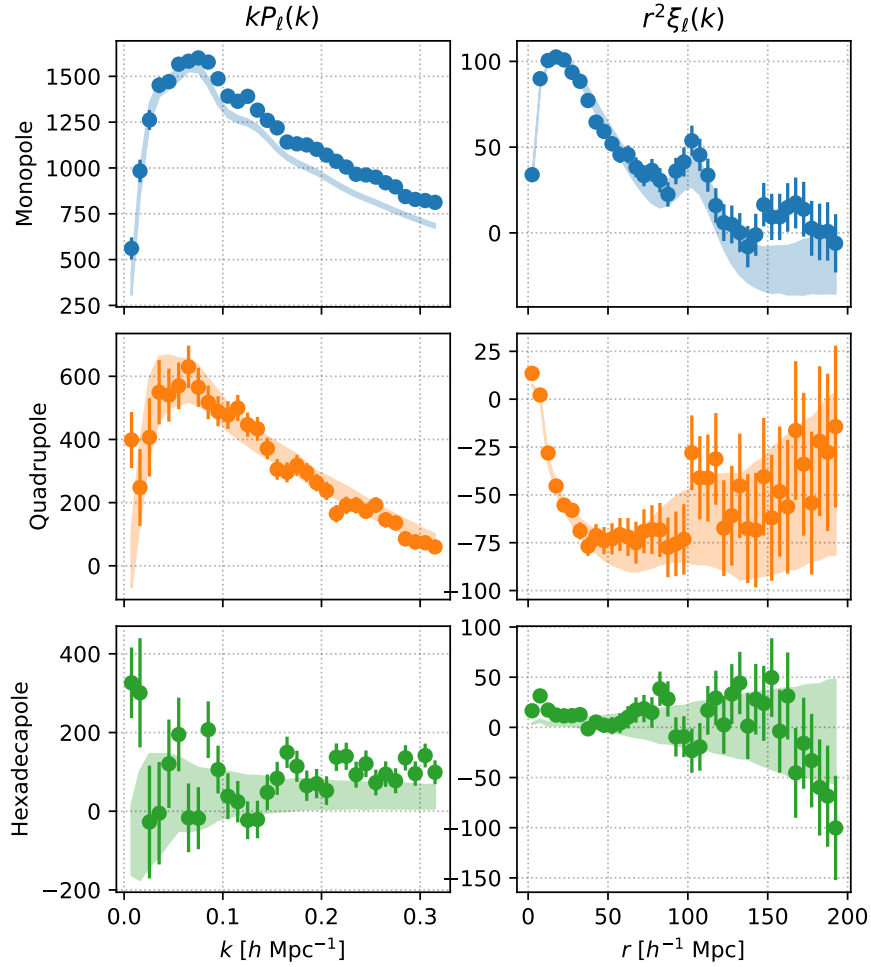


Figure 4.5: Pre-reconstruction multipoles of the two-point statistics for the combination of BOSS DR12 and eBOSS DR16 luminous red galaxies. Left panels contain the power spectrum multipoles (scaled by  $k$ ) and right panels the correlation function multipoles (scaled by  $r^2$ ). Points with error bars are real data while shaded regions show the mean and standard deviation of 1000 realisations of EZMOCKS.

#### 4.4.3 The clustering of eBOSS DR16 LRGs

Figure 4.5 and 4.6 show the power spectrum  $P_\ell(k)$  and correlation function  $\xi_\ell(r)$  multipoles for  $\ell = 0, 2, 4$  for the combined sample of BOSS DR12 and eBOSS DR16 luminous red galaxies and their respective EZMOCK catalogues. Figure 4.5 shows results from standard catalogues while Figure 4.6 displays results from reconstructed catalogues.

The uncertainties on our measurements of  $P_\ell$  and  $\xi_\ell$  were derived by measuring the same statistics on the full sample of 1000 realisations of EZMOCKS. Mock based covariance are computationally expensive but they account for all sources of uncertainty: cosmic variance, shot-noise, and systematic effects. By reproducing with high fidelity the observational features and systematic effects, the covariance matrix obtained with mocks should account for the scatter induced by them. However on small scales the clustering of the mocks is not accurate compared to real data or n-body simulations. Figure 4.7 shows the resulting covariance matrices normalised by the diagonal elements (correlation coefficients). We see how

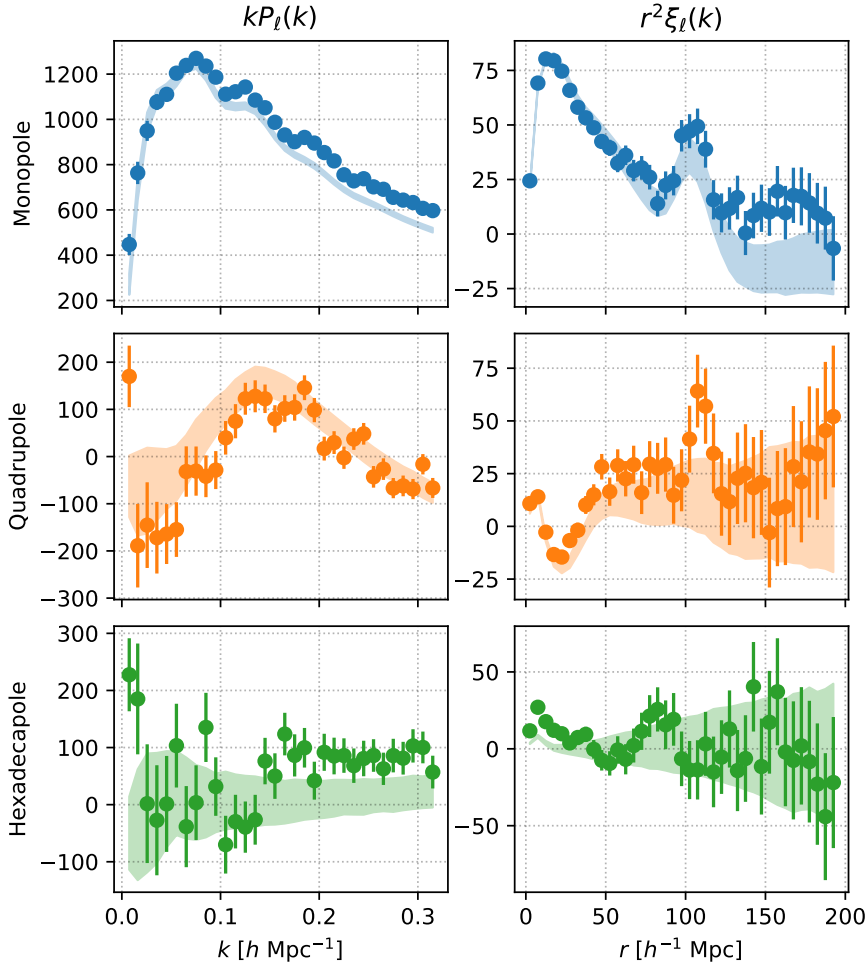


Figure 4.6: Same as Figure 4.5 but for reconstructed catalogues. We used the Zeldovich reconstruction method with removal of redshift-space distortions. The BAO features are more prominent and the quadrupole is significantly reduced compared to results with the standard catalogues.

power spectra multipoles are generally less correlated than correlation function multipoles. Reconstruction also helps reducing these correlations. Also shown are the cross-covariance between Fourier space and configuration space measurements, which will be discussed in section 4.7. As expected, these two spaces are highly correlated since they originate from the same sample of galaxies and randoms.

## 4.5 Baryon acoustic oscillations

The feature from baryon acoustic oscillations (BAO) is clearly visible as oscillations in the power spectrum multipoles and as a sharp peak in the correlation function multipoles (Figures 4.5 and 4.6). In this section I describe how we measure the BAO scale, which is used as a standard ruler in the study of the expansion of the Universe. The work is detailed further in Bautista et al. 2018; Bautista et al. 2021; Gil-Marín et al. 2020.

The model used to fit BAO is based on the linear matter power spectrum with a simpli-

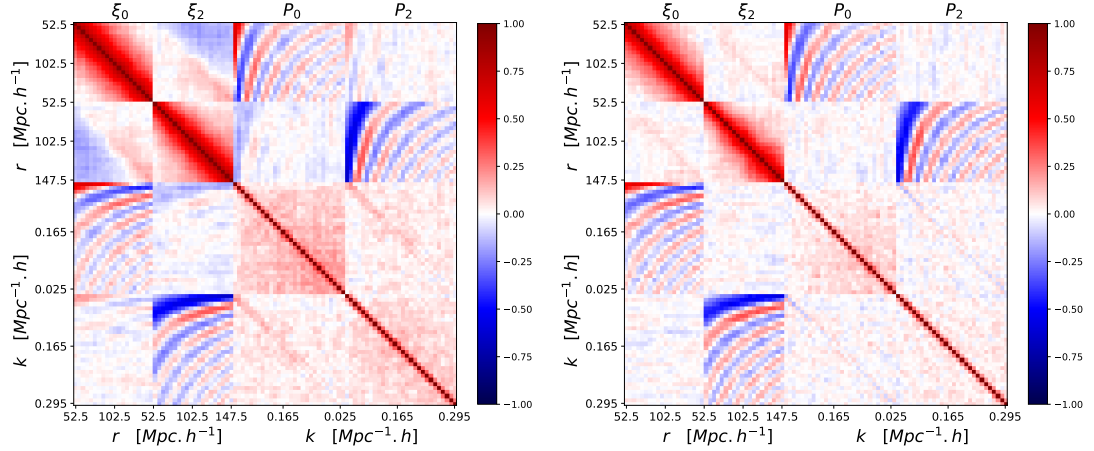


Figure 4.7: Correlation coefficients from covariance matrices obtained from 1000 measurements of power spectrum and correlation function multipoles,  $P_\ell$  and  $\xi_\ell$  respectively, from EZMOCKS reproducing the combined BOSS DR12 and eBOSS DR16 LRG sample. Left panel shows results from the standard catalogues while the right panel shows the results from Zelovich reconstructed catalogues. Figure extracted from Dumerchat and Bautista 2022.

fied treatment of bias, redshift-space distortions and non-linearities. It is basically the same model as used in the BAO measurements from Ly $\alpha$  forests described in section 3.2.4 but without the treatment of metal absorption, high-column density systems and distortions which do not apply here. A linear bias and linear redshift-space distortions is assumed. The non-linearities that broaden the BAO peak are modelled by two Gaussian smoothing terms, one for radial separations and one for transverse separations. We also decompose the peak part from the smooth part of the model, scaling only the peak part by the dilation factors  $\alpha_{\parallel}$  and  $\alpha_{\perp}$ , defined in Eqs. 3.11 and 3.12. The fit is performed in the space of multipoles instead of fitting for the full two-point statistics. Smooth power-laws of separation are added with free amplitudes in order to marginalise any potential information coming from the smooth part of the correlation function. To account for the removal of redshift-space distortions after reconstruction, an empirical damping term is added to the model, such that the linear RSD term  $(1 + \beta\mu^2)$  becomes  $(1 + S(k)\beta\mu^2)$  where  $S(k) \equiv 1 - e^{-k^2\Sigma_{\text{rec}}^2/2}$  (Seo et al. 2016).

The methodology was tested with the sample of 1000 EZMOCKS and also an additional set of 84 NSERIES n-body simulations (section 4.3). In principle, BAO is a feature on large enough scales such that testing the method on approximate mocks is sufficient, although we also looked at results on more realistic n-body simulations. In order to understand potential biases in best-fit BAO parameters, we fit for BAO on the average clustering of all available realisations for EZMOCKS and NSERIES mocks, effectively increasing the volume of the survey by a factor of 1000 and 84 respectively. Thanks to the large number of EZMOCKS, we were able to study the performance of our fitter statistically, by looking at the ensemble of 1000 individual BAO fits.

Figure 4.8 displays the results of fits to the mean clustering of mocks for catalogues pre- and post-reconstruction, also as a function of the fiducial choice for  $\Omega_m^{\text{fid}}$  used in converting distances to redshifts. First we notice how biases are much smaller for reconstructed cat-

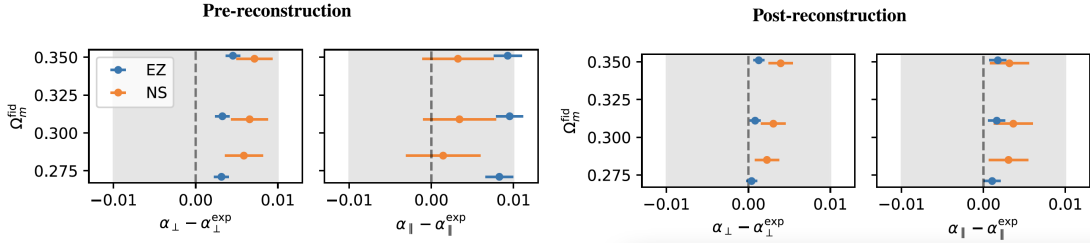


Figure 4.8: Impact of choice of fiducial cosmology in the recovered values of  $\alpha_{\parallel}$  and  $\alpha_{\perp}$  from BAO fits to the stacks of 1000 multipoles the EZMOCKS (blue) and 84 NSERIES mocks (orange), for pre- (left panels) and post-reconstruction (right panels). Associated error bars correspond to the uncertainty estimated from the mean of the mocks. The grey shaded areas correspond to 1 per cent shifts. For comparison, the uncertainty on real data is 1.9 per cent for  $\alpha_{\perp}$  and 2.6 per cent for  $\alpha_{\parallel}$  in the post-reconstruction case.

alogues than standard ones, likely thanks to a better modelling of non-linearities affecting the BAO peak which are reduced in the reconstructed case. Second, uncertainties of reconstructed results are reduced relative to pre-reconstruction catalogues, showing the good performance of the method to remove bulk motions. Third, there is a very weak dependency of our results with the choice of  $\Omega_m^{\text{fid}}$ , which is reassuring.

Another important test of our methodology consists in checking our uncertainty estimation. With the sample of 1000 BAO measurements from mocks, we can study them statistically. Our uncertainties are derived by scanning the likelihood profile and finding the intervals where the parameters  $\alpha_{\perp}$  and  $\alpha_{\parallel}$  increase the  $\chi^2$  by unity<sup>8</sup> relative to its minimum. First, we compare the standard deviation of all best-fit values to the average uncertainty found in all 1000 mocks. If uncertainties are correctly estimated, then these two numbers should match. Second, we could look at the pull distribution of each parameter  $x_i$ , defined as  $Z_i \equiv (x_i - \bar{x}_i)/\sigma_{x_i}$ . If the uncertainties are well behaved and Gaussian, then we should obtain  $\bar{Z}_i = 0$  and  $\text{Var}(Z_i) = 1$ . Table 6 in Bautista et al. 2021 shows these statistics for BAO measurements. The results validate our methodology and we applied it to data.

Figure 4.9 shows the best-fit correlation function model and the BAO constraints for the eBOSS DR16 LRG sample. Converting the constraints in  $(\alpha_{\perp}, \alpha_{\parallel})$  into ratios of distances, we obtained:

$$\mathbf{D}_{\text{BAO}, \xi_{\ell}} = \begin{pmatrix} D_M/r_d \\ D_H/r_d \end{pmatrix} = \begin{pmatrix} 17.86 \pm 0.33 \\ 19.34 \pm 0.54 \end{pmatrix}, \quad (4.8)$$

which correspond to a BAO measurement at 1.9 per cent in the transverse direction and 2.8 per cent in the radial direction. Radial and transverse BAO measurements are correlated by -33% for this dataset. These are the best constraints obtained from galaxies at redshifts above 0.6.

## 4.6 Redshift-space distortions

The analysis of redshift-space distortions (RSD) aims to measure the growth-rate of structures  $f$ , as introduced in section 1.5.6, by quantifying the anisotropies in the clustering of

<sup>8</sup>In the case marginalised errors in one parameter. For two parameters, the 68 and 95% confidence intervals are defined by  $\Delta\chi^2 = 2.3, 5.9$  respectively.



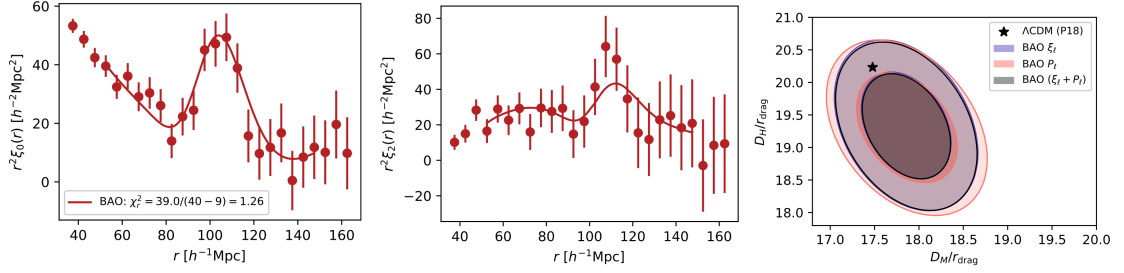


Figure 4.9: Two left panels show the best-fit model for the correlation function monopole and quadrupole for the reconstructed eBOSS LRG sample. The right panel shows the two-dimensional constraints on the distance to sound horizon ratios, for the configuration and Fourier space BAO analyses. The consensus BAO result is shown in grey. The star points to the prediction of a flat  $\Lambda\text{CDM}$  model with Planck 2018 best-fit parameters.

galaxies. As we seen before, assuming an incorrect fiducial cosmology when converting redshifts into distances also introduces anisotropies (the AP effect). Therefore in practice, an RSD analysis yields simultaneously the combination  $f(z)\sigma_8(z)$  and the distance ratios  $D_H(z)/r_{\text{drag}}$ ,  $D_M(z)/r_{\text{drag}}$ .

The main differences with respect to the BAO analysis (see previous section) is that the information about anisotropies come from the full shape of the two-point statistics, not only the peak. This complicates the RSD analysis in that we need accurate models for the full shape of the clustering down to some scale, and we cannot add empirical functions of scale to be marginalised over as we did in the BAO analysis. Due to non-linearities of matter clustering and galaxy bias, it is increasingly hard to model scales below those well described by linear theory. For these mildly non-linear scales, we need more advanced perturbation theory models. A huge variety of theoretical models have been developed in the past years (see Bernardeau et al. 2002 for a review).

In the analysis of eBOSS data, we used two types of models: the first one is based on regularised perturbation theory with second order bias and RSD correction terms (Taruya et al. 2010; Taruya et al. 2012). We commonly refer to it as the *TNS model*. The second one is based on convolution Lagrangian perturbation theory with Gaussian streaming (Carlson et al. 2013; Wang et al. 2014). We refer to it as the *CLPT-GS model*. I implemented these models in python language in the package BAOPY<sup>9</sup> though this version was not the one used in the analysis.

The TNS model is expressed as:

$$P_g(k, \mu_k) = D_{\text{FoG}}(k, \mu_k) \left[ P_{g,\delta\delta}(k) + 2f\mu_k^2 P_{g,\delta\theta}(k) + f^2\mu_k^4 P_{\theta\theta}(k) + b_1^3 A(k, \mu_k, \beta) + b_1^4 B(k, \mu_k, \beta) \right], \quad (4.9)$$

where  $b_1$  is the first order linear galaxy bias,  $\beta \equiv f/b_1$ , ( $P_{g,\delta\delta}, P_{g,\delta\theta}, P_{\theta\theta}$ ) are the non-linear isotropic galaxy auto power spectrum, galaxy-velocity divergence cross power spectrum and the velocity divergence auto power spectrum, respectively, which are themselves functions of first and second order bias parameters. These functions are provided by the RegPT scheme, implemented in PYREGPT<sup>10</sup>. The functions  $A$  and  $B$  are the RSD correction terms. The function  $D_{\text{FoG}}$  can be either a Gaussian or a Lorentzian function of  $k_{\parallel} = k\mu_k$ , with a characteristic scale  $\sigma_{\text{FoG}}$  corresponding to the velocity dispersion on small scales. This model is computed in Fourier space, but it can be Fourier transformed to obtain a model in configuration space.

<sup>9</sup><https://baopy.readthedocs.io/>

<sup>10</sup><https://github.com/adematti/pyregpt>

The CLPT-GS model is uniquely defined in configuration space, and writes as:

$$1 + \xi_{g,s}(r_{\perp}, r_{\parallel}) = \int \frac{1}{\sqrt{2\pi} [\sigma_{12}^2(r) + \sigma_{\text{FoG}}^2]} [1 + \xi_g(r)] \exp - \frac{[r_{\parallel} - y - \mu v_{12}(r)]^2}{2 [\sigma_{12}^2(r) + \sigma_{\text{FoG}}^2]} dy \quad (4.10)$$

where  $\xi_g(r)$ ,  $v_{12}(r)$ , and  $\sigma_{12}(r)$  are obtained from CLPT and contains two bias terms. The dependency with  $f$  is via a scaling of the functions  $v_{12}, \sigma_{12}$  that are proportional to velocity. The integral corresponds to the convolution by velocities in the radial direction, characteristic of the streaming formalism (see Kuruvilla and Porciani 2018 and references therein for an introduction of the streaming model).

Both the TNS and CLPT-GS models had to be tested against n-body simulations. We were particularly interested in the scales of validity of these models for the case of the eBOSS LRG sample. Given the number of available realisations, we used the n-body suite of 84 realisations of NSERIES simulations produced for the BOSS CMASS sample. We fitted the average correlation function of these simulations and measured the parameters  $f\sigma_8, \alpha_{\perp}, \alpha_{\parallel}$ . Figure 4.10 presents the results compared to expected values for both TNS and CLPT-GS models. The shaded green area shows the fiducial choice to be used in the data.

Using the same simulations, we tested the dependency of our constraints to the choice of fiducial cosmology. Note that the fiducial cosmology enters twice in the modelling, first when converting redshifts to distances; second, when computing the linear matter power spectrum template  $P_m^{\text{lin}}(k)$ , used as main ingredient for both models. Figure 4.11 presents the results for three choices of fiducial cosmology. While for  $\alpha_{\perp}, \alpha_{\parallel}$  the biases or trends are negligible,  $f\sigma_8$  shows a strong dependency with  $\Omega_m^{\text{fid}}$ , shown as crosses in the figure. This dependency is mainly caused by the assumed  $\sigma_8$  value, that corresponds to the one derived from the fiducial  $P_m^{\text{lin}}(k)$  via Eq. 1.18. The resulting  $f$  of the fit is scaled by this  $\sigma_8$ . We can reduce this dependency by recomputing  $\sigma_8$  using  $R = 8\alpha h^{-1}\text{Mpc}$ , where  $\alpha \equiv \alpha_{\perp}^{(2/3)}\alpha_{\parallel}^{(1/3)}$  is the isotropic dilation factor obtained from the best-fit  $\alpha_{\perp}, \alpha_{\parallel}$ . In effect, this keeps the scale at which  $\sigma_8$  is fitted fixed relative to the data in units of  $h^{-1}\text{Mpc}$ , which only depends on  $\Omega_m^{\text{fid}}$ . This is an alternative approach to the recently proposed  $\sigma_{12}$  parametrisation (Sánchez 2020), where the radius of the top-hat function is set to  $R = 12\text{Mpc}$  (in units of Mpc) instead of  $R = 8h^{-1}\text{Mpc}$  (in units of  $h^{-1}\text{Mpc}$ ). The circles in the right panel of Figure 4.11 shows the results with the recomputed  $\sigma_8$  for each case, showing that this procedure indeed reduces the trend of  $f\sigma_8$  with  $\Omega_m^{\text{fid}}$ . This topic is also discussed in Appendix A of Alam et al. 2021.

After validating the model, we tested the impact of different halo occupation distribution (HOD) models and observational effects, both introduced in simulations. We found that none of these modifications introduces significant biases our constraints of  $(f\sigma_8, \alpha_{\perp}, \alpha_{\parallel})$ , though the shifts we observe were added as systematic uncertainties in quadrature to the statistical ones. Table 10 of Bautista et al. 2021 summarises the systematic errors which are assumed to be diagonal (no covariance between the three parameters) for simplicity. For our RSD analysis, systematic errors are about half of the statistical ones for  $f\sigma_8$  and slightly less for  $\alpha_{\perp}, \alpha_{\parallel}$ . In quadrature, final uncertainties increase by up to 20% due to systematic errors.

We also attempted to analyse our biases and uncertainties statistically with the set of 1000 EZMOCKS. However, an important caveat is that these mocks only solve approximatively the clustering, and therefore are not the best suit to test RSD fitting. We proceeded to fit the 1000 realisations and we compared the average uncertainty to the standard deviation of best-fit parameters among all mocks. These numbers agree sufficiently well, particularly for  $f\sigma_8$ . The pull distributions of each parameters do not point to any large under/over estimation of uncertainties.

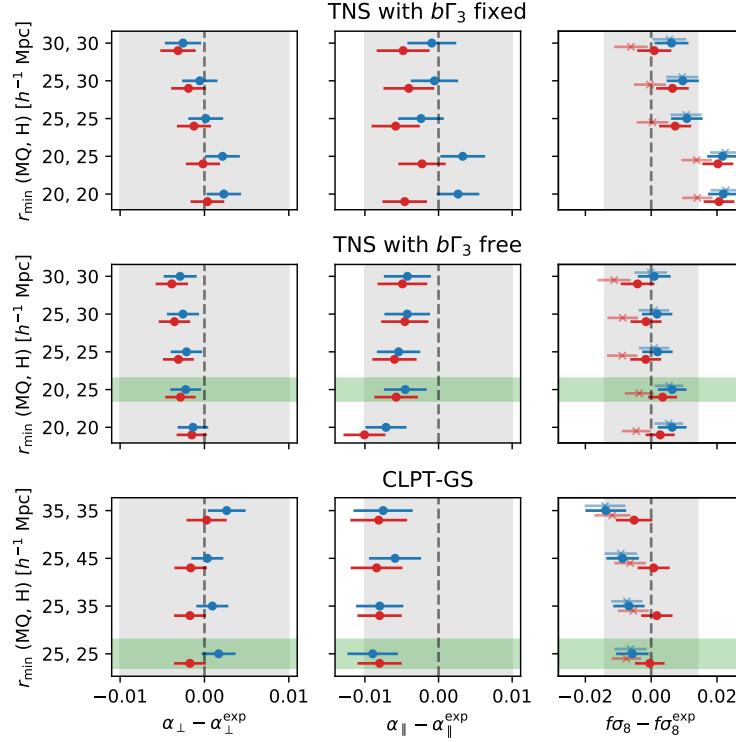


Figure 4.10: Biases in the measurement of  $f\sigma_8, \alpha_{\parallel}, \alpha_{\perp}$  obtained from full-shape fits to the average of 84 multipoles from the NSERIES mocks as a function of the separation range used. The y-axis displays the value of the minimal separation  $r_{\min}$  used in fits of the monopole, quadrupole (MQ) and hexadecapole (H). Top and mid rows display results for the TNS model when fixing or letting free the bias parameter  $b\Gamma_3$  respectively. Bottom row presents results for the CLPT-GS model. The blue circles correspond to the analysis using  $\Omega_m^{\text{fid}} = 0.286$  (the true value of simulations) while the red circles correspond to  $\Omega_m^{\text{fid}} = 0.31$ . The crosses in the  $f\sigma_8$  panels correspond to their values before the scaling of  $\sigma_8$  discussed in Figure 4.11. The gray shaded areas correspond to 1 per cent errors in  $\alpha_{\perp}, \alpha_{\parallel}$  and to 3 per cent in  $f\sigma_8$ . The green shaded area shows our choice for baseline analysis for TNS and CLPT-GS models.

Figure 4.12 presents the best-fit TNS and CLPT-GS models to the correlation function multipoles of the eBOSS DR16 LRG sample. Both models are basically indistinguishable even though they are derived differently. The constraints are also similar between the two models, so we combined them into a single results by averaging their best-fit parameters and their covariances. Figure 4.13 shows the final constraints on  $f\sigma_8$  and the distance ratios, compared to the same analysis in Fourier space. We also show the consensus result computed using the method by Sánchez et al. 2017, which assumes Gaussianity of each posterior distribution (see the next section for a further discussion). The final RSD constraints from the correlation function fits are:

$$\mathbf{D}_{\text{RSD}, \xi_\ell} = \begin{pmatrix} D_M/r_d \\ D_H/r_d \\ f\sigma_8 \end{pmatrix} = \begin{pmatrix} 17.42 \pm 0.40 \\ 20.46 \pm 0.70 \\ 0.460 \pm 0.050 \end{pmatrix} \quad (4.11)$$

which correspond to a 11% measurement of  $f\sigma_8$  at  $z_{\text{eff}} = 0.7$ . This uncertainty is slightly reduced when combining with the Fourier space results and with BAO results from the previous

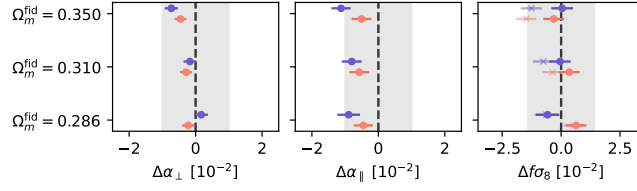


Figure 4.11: Biases in best-fit parameters for both CLPT-GS (blue) and TNS (red) models from fits to the average multipoles of 84 NSERIES mocks. Shaded grey areas show the equivalent of 1 per cent error for  $\alpha_{\perp}$ ,  $\alpha_{\parallel}$  and 3 per cent for  $f\sigma_8$ . In the right panel, crosses indicate  $f\sigma_8$  values when  $\sigma_8$  is not recomputed as described in the text. The true cosmology of the mocks is  $\Omega_m = 0.286$ . For reference, the errors on our data sample are  $\sim 2, 3$  and  $10$  per cent for  $\alpha_{\perp}$ ,  $\alpha_{\parallel}$ ,  $f\sigma_8$  respectively.

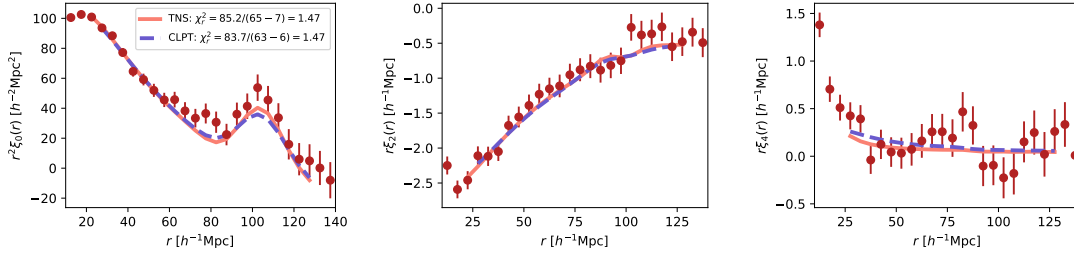


Figure 4.12: Best-fit models of redshift-space distortions (RSD) to the eBOSS DR16 LRG multipoles. Left, mid and right panel display mono, quad and hexadecapole, respectively. Note that only the monopole is scaled by  $r^2$  while the others are scaled by  $r$ . The CLPT-GS model is shown by the blue dashed line while the TNS model is shown by the red solid line.

section. The combination between Fourier and configuration space results will be discussed in the following section.

## 4.7 Joint clustering analysis in Fourier and Configuration space

From the same catalogue of angular positions and redshifts, we can perform two different clustering analyses: one in Fourier space (FS) and the other in Configuration space (CS). How to define a consensus result from both analysis which might yield slightly different results? In this section I expose the work lead by my PhD students Tyann Dumerchat and Vincenzo Aronica on how to combine BAO and RSD constraints from both spaces into a single consensus result.

Sánchez et al. 2017 proposed a method to combine two correlated Gaussian posteriors into a single one. I refer to it as the Gaussian approximation (GA) method. The GA method was used in BOSS (Alam et al. 2017) and it was slightly improved in eBOSS (Bautista et al. 2021). The main idea is that we want to combine a set of vectors containing the best-fit parameters of different analysis (and their respective error matrix) into a single vector of parameters and a single error matrix. For example in BOSS and eBOSS, we aimed at combining the vectors  $\vec{p} \equiv f\sigma_8, \alpha_{\perp}, \alpha_{\parallel}$  obtained from FS and CS analyses. The GA method uses an estimate the correlations between  $\vec{p}_{\text{FS}}$  and  $\vec{p}_{\text{CS}}$ , and built a large covariance matrix used to combine the vectors and error matrices. Currently the best method to obtain these correlations be-

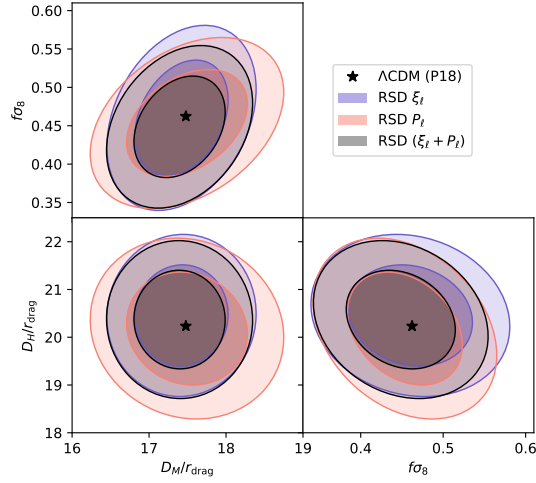


Figure 4.13: Constraints on  $D_M/r_d$ ,  $D_H/r_d$  and  $f\sigma_8$  at  $z_{\text{eff}} = 0.7$  from the full-shape RSD analysis of the completed eBOSS LRG sample. Contours show 68 and 95 per cent confidence regions for the analyses in configuration space (blue), Fourier space (red) and the combined (grey). The expected values in a flat  $\Lambda$ CDM model with best-fit parameters from Planck 2018 results is indicated as a black star.

tween FS and CS results is using mock catalogues. This provides with an estimate of a sort of “ensemble averaged” correlations (though limited by the number of realisations). However, the real data is a single realisation of the ensemble, and the correlations between FS and CS are also susceptible to fluctuate, just as the individual parameters and their uncertainties. In Bautista et al. 2021 we developed a method to account for this fact, where we adjust the correlations for each realisation based on the vectors  $\vec{p}$  and their error matrices of that particular realisation.

In Dumerchat and Bautista 2022, we developed an analysis where we fit jointly the correlation function and power spectrum while accounting for their covariance. I refer to it as the joint space (JS) method. The covariance is built at the multipole level, using a set of mock catalogues. Figure 4.7 shows the correlation coefficients of CS and FS multipoles, including the cross-correlation terms. The wiggles observed in the cross-correlation terms can be simply explained by the fact that CS and FS are related via Hankel transforms (see Appendix A of Dumerchat and Bautista 2022 for an explicit derivation). We tested and validated this framework on the BAO analysis of the eBOSS LRG sample (section 4.5) where we fit for a single  $\alpha_{\perp}$  and  $\alpha_{\parallel}$ , though amplitudes of power-laws are independent for each space. The main advantage of the JS method is that we do not assume Gaussian posteriors at the parameter level, only at the multipole level when fitting our model (which is quite common assumption). The posteriors on the resulting  $(\alpha_{\perp}, \alpha_{\parallel})$  can also be non-Gaussian. Figure 4.14 displays few examples from BAO fits to mock realisations where the CS and FS posteriors are Gaussian (left) or non-Gaussian (right panels), and the resulting combination from GA and JS methods. For the non-Gaussian cases, the GA method fails to correctly describe the posterior, underestimating uncertainties.

Vincenzo Aronica is currently working on the application of joint space analysis for the case of redshift-space distortions (RSD). He implemented the TNS model (section 4.4.3) in Fourier and configuration space in the `BAOPY` package. The tests on the suite of 84 realisations of `NSERIES` n-body simulations are promising, though there are some subtleties to solve

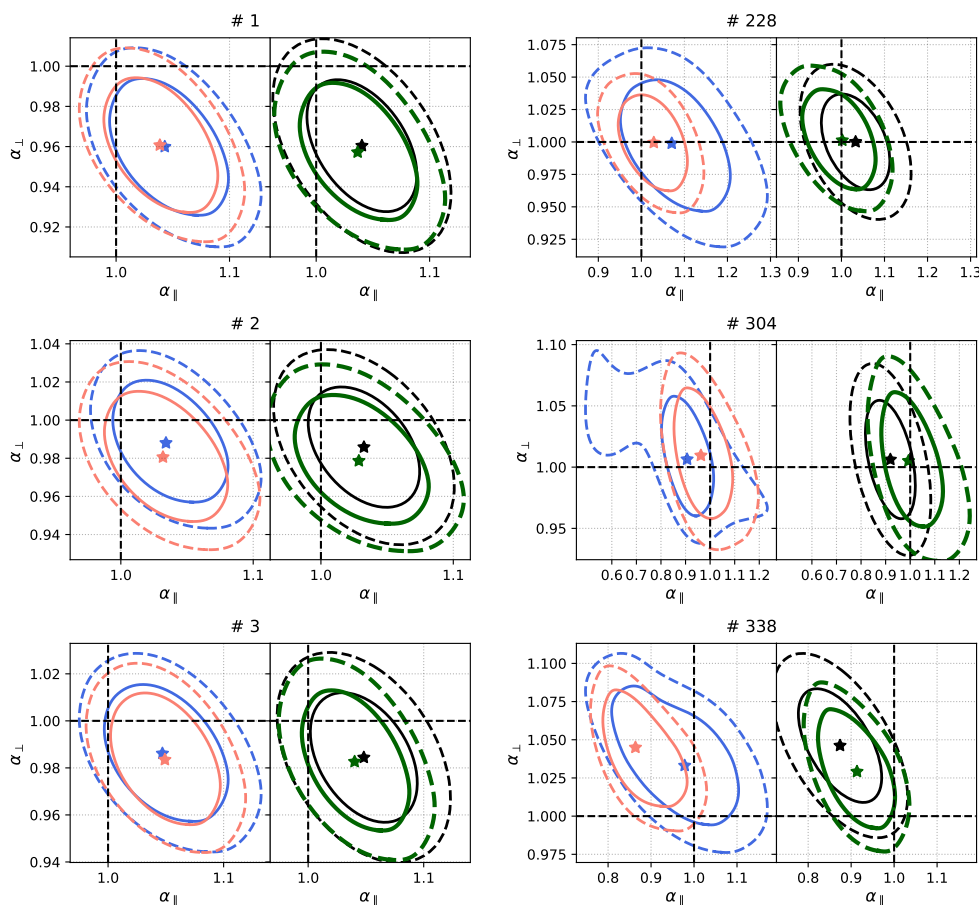


Figure 4.14: Posteriors on the BAO fits of some EZMOCKS realisations. We display cases where contours found in configuration or Fourier space are Gaussian (left) and non Gaussian (right). Red contours are for Fourier space (FS), blue for configuration space (CS), black for the Gaussian approximation (GA) and green for the joint space (JS) fits. The expected value is the intersection of the dotted black lines, and the best fit values are described by a star for the JS and GA methods. We can see how the JS method yields better combined results that are not necessarily Gaussian. Extracted from Dumerchat and Bautista 2022.

related to the relative normalisation (or linear bias) amplitudes between FS and CS. Currently we use the same values of biases for both spaces, in addition to  $\alpha_{\perp}$ ,  $\alpha_{\parallel}$ ,  $f\sigma_8$ . Vincenzo will recompute power spectra for a given set of mock catalogues using a consistent normalisation and window function calculations provided by PYPPOWER.

The work on BAO and RSD joint fits in Fourier and configuration space will be hopefully applied to the analysis of DESI data.

## 4.8 Cross-correlation with radio surveys

The study of large-scale structures will soon be improved thanks to observations of the distribution of neutral hydrogen (HI) via the emission of photons from the 21cm hyperfine transition of the electron. Future 21cm surveys such as the Square Kilometre Array (SKA) will be able to observe the Universe with intensity mapping up to redshifts of 3, opening an alternate

window to that provided by spectroscopic surveys. The term *intensity mapping* refers to the fact that we measure the integrated flux from unresolved sources at a given location in space.

During my postdoctoral stay at the University of Portsmouth, I contributed to the study of a currently available 21cm survey produced by the Green Bank Telescope (GBT). I was particularly interested in the study of cross-correlations between the 21cm observed by GBT and the eBOSS galaxies observed in the same volume. With the cross-correlation function we could infer the product of the average neutral hydrogen fraction  $\Omega_{\text{HI}}$  at redshift  $z_{\text{eff}} = 0.8$  with the bias of HI  $b_{\text{HI}}$ . These parameters are important ingredients for models of clustering of the 21cm emission. I overview the main findings of this work in this section but further details are described in Wolz et al. 2022.

The use of 21cm intensity mapping to measure the large-scale structures has been extensively studied theoretically (Bharadwaj et al. 2001; Battye et al. 2004; McQuinn et al. 2006; Chang et al. 2008) and with n-body simulations (Mesinger et al. 2011; Villaescusa-Navarro et al. 2018 and references therein). On large scales, 21cm flux fluctuations trace those of dark matter density, while on smaller scales the physics of the gas and its connection to halos complexifies the modelling. Including redshift-space distortions, a linear model for the 21cm power spectrum is essentially the same as the linear model for galaxy clustering. In Villaescusa-Navarro et al. 2018, an extensive study using the IllustrisTNG n-body hydro simulation quantified several ingredients commonly used in more complex models of the 21cm power spectra, such as the HI content of halos, the HI probability density function, the HI bias with respect to matter and others.

The data sets we used in our study were the eBOSS luminous red galaxy (LRG) and emission line galaxy (ELG) samples in one hand, and the Green Bank Telescope (GBT) 21cm maps in the other hand. The eBOSS LRG sample is already described in section 4.1 while the ELG sample is described in Raichoor et al. 2020. The GBT maps are observations of a 100 deg<sup>2</sup> patch of sky at frequencies from 700 to 900 MHz divided into 256 channels, corresponding to HI emission between  $0.6 < z < 1.0$ , perfectly overlapping with our galaxy samples in area and redshift. The angular resolution of HI observations varies from 0.31 deg at 700 MHz to 0.25 deg at 900 MHz. In order to reduce the impact of varying PSF, all channels are convolved by an artificial beam of 0.44 deg. The maps we used are an extension to those produced by Masui et al. 2013, including extra scans which increase the area and depth. The raw data is contaminated by radio frequency interference and some channels are therefore masked. Foregrounds are removed with a technique named *FastICA*, which stands for fast independent component analysis. I was provided several maps with different number of foreground components removed. Once the final area of GBT maps was fixed, I cut the eBOSS LRG and ELG samples to the same area. Figure 4.15 shows a redshift slice (or one channel) of the foreground-cleaned GBT HI map and the LRG and ELG angular distribution in the same patch of sky.

The cross-correlations between eBOSS galaxies and 21cm temperature maps was computed in Fourier space by Wolz et al. 2022 and in configuration space by myself. The cross-correlation estimator I used is the same used for the quasar- Ly $\alpha$  forest cross-correlation described in section 3.2.2, which assumes that galaxies (in this case) are shot-noise dominated and that voxels of the 21cm map are independent. These assumptions are not quite valid for our sample but I left further developments for future work. For instance, one could perform the study of cross-correlations with a maximum likelihood method (see e.g. section 5.3.1) or in Fourier space but correctly accounting for the convolution by the window function of these surveys, which was not explicitly done in Wolz et al. 2022.

Figure 4.16 displays the monopoles and quadrupoles of the cross-correlation function between 21cm temperature fluctuations and eBOSS galaxies. Both ELG and LRG show non-



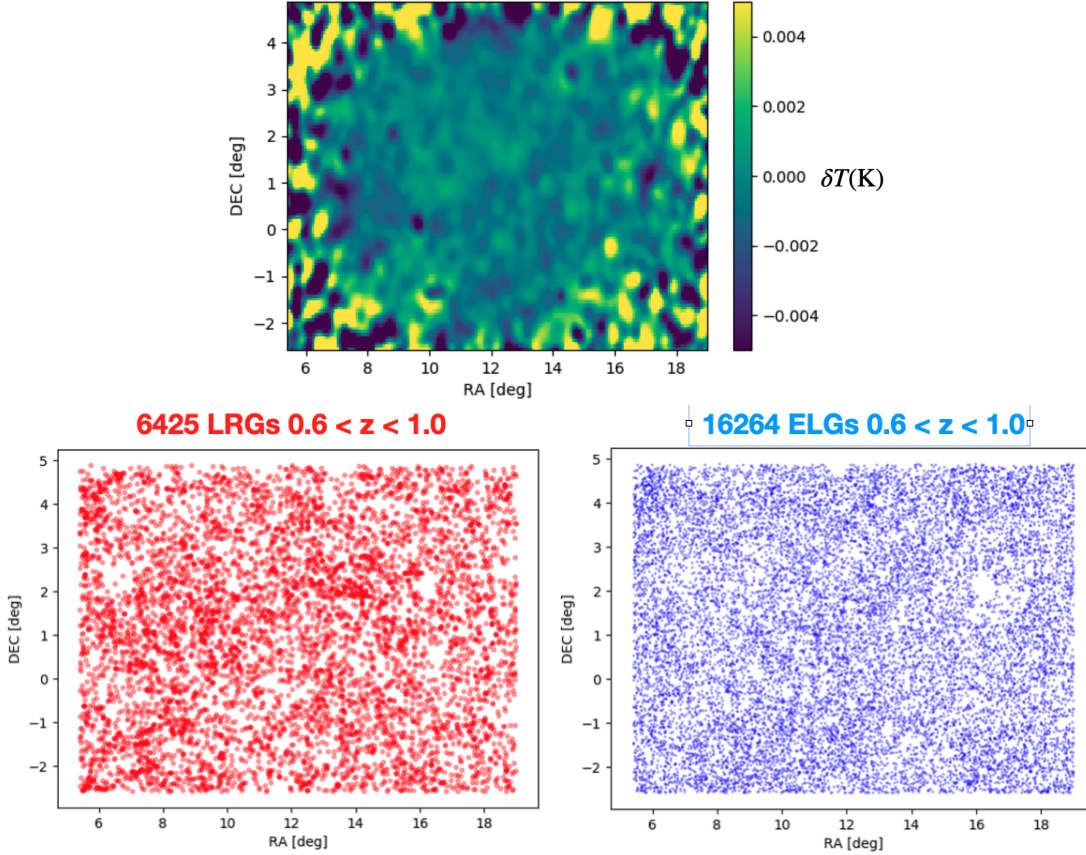


Figure 4.15: Top panel: a redshift slice of the 21cm HI map from the Green Bank Telescope after foreground removal. Bottom panels: the angular distribution of all eBOSS LRGs and ELGs falling in the same sky area.

zero correlations with the temperature maps that extend to cosmological scales. In particular, we clearly see a non-zero quadrupole, which was measured for the first time in this analysis. The covariance matrix of the multipoles was computed with the jackknife technique, using 96 independent subsets of the data. Due to the estimator I used, neighbouring separation bins in the monopole are strongly correlated.

I attempted to fit the temperature-galaxy cross-correlation with the following model:

$$P(k, \mu) = \frac{b_{\text{gal}} b_{\text{HI}} \bar{T}_{\text{HI}}}{(1 + k^2 \mu^2 \Sigma_{\text{FoG}}^2)^2} (1 + \beta_{\text{gal}} \mu^2) (1 + \beta_{\text{HI}} \mu^2) P_{\text{lin}}(k) e^{-k^2 R_{\text{beam}}^2 (1 - \mu^2)/2} \quad (4.12)$$

which includes a term for Fingers-of-God with dispersion  $\Sigma_{\text{FoG}}$ , and a damping in the transverse (angular) direction to account for the effect of the beam of the GBT instrument, with characteristic scale  $R_{\text{beam}}$ . The solid lines in Figure 4.16 show the best-fit models, which have in total 6 free parameters. The  $\chi^2$  values are also shown in the figure, which correspond to 30 data points (24 degrees of freedom). The fits are generally well behaved. I do not show values for the best-fit parameters and their uncertainties because of the caveats I mentioned before related to the estimator and the covariance matrices. Also at the time of my analysis, we did not have mock catalogues to test my methodology.

Wolz et al. 2022 proceeded with the analysis of the cross power spectrum monopole and fitted for the amplitude of the signal, after testing our methods on mock catalogues. Us-



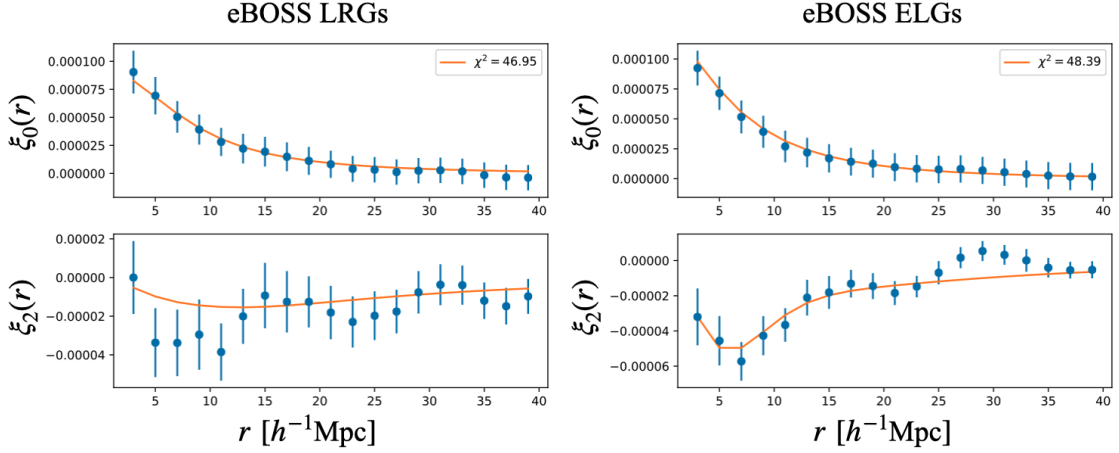


Figure 4.16: Cross-correlation function multipoles between 21cm temperature fluctuation maps from the Green Bank Telescope and the eBOSS samples of Luminous Red Galaxies (left) and Emission Line Galaxies (right). Uncertainties are computed with jackknife resampling and solid lines correspond to the best-fit models.

ing different scale ranges, we obtained constraints on the combination  $\Omega_{\text{HI}} b_{\text{HI}} r$  at  $z_{\text{eff}} = 0.8$ , where  $r$  is an empirical cross-correlation term that depends on the galaxy sample. Table 1 of Wolz et al. 2022 summarises results. For instance, for conservative choices for the analysis, we find  $\Omega_{\text{HI}} b_{\text{HI}} r = (0.48 \pm 0.12) \times 10^{-3}$  for GBT x ELGs and  $(0.38 \pm 0.12) \times 10^{-3}$  for GBT x LRGs, corresponding to 5 and  $4.2\sigma$  detections respectively.

This work is on the pathway for precision measurements with future 21cm surveys such as SKA (Square Kilometre Array Cosmology Science Working Group et al. 2020).

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## Chapter 5

# The low-redshift Universe and its velocities

At low redshifts, peculiar velocities of galaxies can be directly measured and be used to learn about gravity on large scales. As presented in chapters 1 and 4, velocities are generally not observable other than by their indirect impact on the galaxy clustering through redshift-space distortions (RSD). At  $z < 0.2$ , the volume that can be probed by galaxy surveys is insufficient to obtain precise measurements of the growth-rate of structures  $f$  solely with RSD. The additional information from peculiar velocities significantly improves these constraints, to the point where one can perform consistency tests of the standard cosmological model or of our current model for gravity: general relativity.

In this chapter I will expose the recent work we developed within the cosmology group at the Centre de Physique des Particules de Marseille in trying to combine galaxy surveys and peculiar velocities, particularly those obtained from type-Ia supernovae. Most of the work was carried out by the PhD student Bastien Carreres whom I co-supervise.

### 5.1 Measuring peculiar velocities

As discussed in section 1.5.6, redshifts of galaxies can be seen as a mixture of two distinct types of redshift. The first one is due to the expansion of the Universe and we refer to it as *cosmological redshift*  $z_{\text{cos}}$ . The second one is due to Doppler effect and the peculiar velocities of galaxies with respect to a comoving frame, we refer to it as *peculiar redshift*  $z_{\text{pec}}$ . The total observed redshift  $z_{\text{obs}}$  can be written as:

$$1 + z_{\text{obs}} = (1 + z_{\text{cos}})(1 + z_{\text{pec}}) \quad (5.1)$$

Measuring peculiar velocities or redshifts  $z_{\text{pec}}$  means breaking the degeneracy between  $z_{\text{cos}}$  and  $z_{\text{obs}}$ . The observed redshift of a given galaxy can be precisely measured via spectroscopy but its cosmological redshift requires a measurement of its distance or age, of which we can derive  $z_{\text{cos}}$  through a cosmological model. Therefore, peculiar velocity surveys can be seen simply as distance surveys.

Measuring distance of individual objects on cosmological scales is quite challenging since it commonly involves standard or standardisable candles. For cosmological distances, three types of standardisable candles have been used by peculiar velocity surveys: spiral galaxies, elliptical galaxies and type-Ia supernovae. Many decades of studies of galaxies have pointed to interesting correlations between their luminosity or absolute magnitudes, which are not



directly observable, to properties or physical quantities that can be directly observed and do not depend on the galaxy distance. Two famous correlations are the Tully-Fisher (TF) relation for spiral galaxies and the Fundamental Plane (FP) for ellipticals, which we will describe shortly after. Needless to say, type-Ia supernovae (SNIa) have been proven to be excellent standardisable candles when accounting for correlations between their luminosity and the duration and colour of their light-curves. As discussed in section 1.5.2, SNIa are by themselves a key cosmological probe of the Universe's expansion.

The precision of distance measurements is highly dependent on the modelling of correlations between luminosity, from which we derive luminosity distances, and their observed properties. Even after standardisation, each method remains with some intrinsic scatter in distances. For instance, TF and FP distances have generally a scatter of about 20% while SNIa can reduce it to about 8%. For cosmological applications, one SNIa distance is worth 4 to 5 TF or FP distances to obtain the same precision.

### 5.1.1 The Tully-Fisher relation of spiral galaxies

The Tully-Fisher relation is an empirical link between the mass of spiral galaxies and their rotational velocity at the regime where it does not vary much with distance to the center. Tully and Fisher 1977 were the first to find a correlation between a distance-independent observable, the asymptotic rotational velocity, and mass or luminosity, and to use it as a method to infer galaxy distances. The physical origin of the relation is the virial theorem of a collapsed object.

Later works found that the infrared luminosity correlates better with rotational velocity, i.e., it reduces the remaining scatter. This is because infrared wavelengths are less sensitive to dust extinction than optical wavelengths. McGaugh et al. 2000 have later shown that the standard Tully-Fisher relation breaks down for galaxies with small rotation velocities. However, those galaxies contain relatively more gas than stars. If the total baryonic mass is used instead of only stellar mass, the relation holds even for these smaller galaxies. This is known today as the baryonic Tully-Fisher (BTF) relation.

One of the largest samples of BTF distance measurements to date is the one provided by the fourth version of the Cosmicflows project (Kourkchi et al. 2022), containing about 10 000 distances. Figure 5.1 and 5.2 show the angular and redshift distribution of this catalogue. The BTF relation of the sample is calibrated using a set of 600 galaxies in 20 well known clusters. They used information from HI fluxes from several radio telescopes. The HI fluxes are better tracer of the gas content of the galaxies and the HI line width is related to the rotation velocity of these galaxies. For stellar masses, they used photometry from SDSS in optical and the WISE satellite in the infrared. The inferred scatter in their TF relation is of about 45% in the distance modulus (or absolute magnitudes) which corresponds to a scatter of  $\ln(10)/5 \times 45\% = 18\%$  in distances. Their sample extends up to  $z \sim 0.05$ . Hoffman et al. 2021 describes how to compute peculiar velocities from this sample of TF distances, including a correct treatment of uncertainties.

In section 5.3 I will present some of the growth-rate measurements using this or a previous version of the CosmicFlows catalogue of distances.

### 5.1.2 The Fundamental Plane of elliptical galaxies

The equivalent of the Tully-Fisher relation for elliptical galaxies is named the Fundamental Plane (FP, Djorgovski and Davis 1987). It is an extension in two-dimensions of a previously thought-to-be one-dimensional relation between the luminosity of the elliptical galaxy

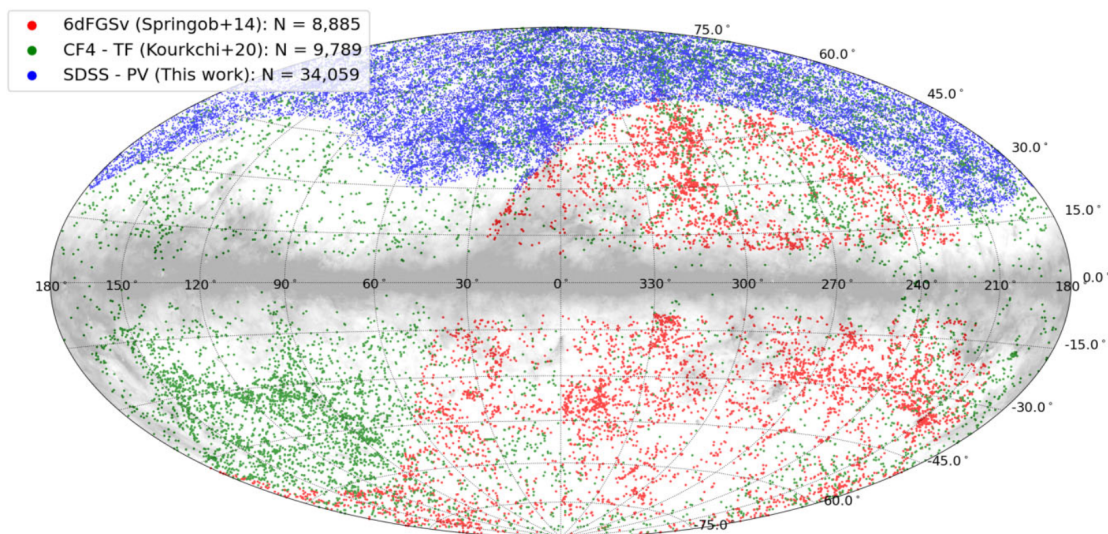


Figure 5.1: Angular distribution of peculiar velocity samples in Galactic coordinates: SDSS (blue) and 6dFGSv (red) from Fundamental Plane and Cosmicflows-4 (green) from Tully-Fisher relation. Figure extracted from Howlett et al. 2022.

and the velocity dispersion of their stars near their centres. The one-dimensional version is known as the Faber-Jackson relation (Faber and Jackson 1976). Later, this relation was extended to two-dimensions to relate the effective radius of the galaxy (in kpc) to the velocity dispersion and the average surface brightness within the effective angular radius. The last two quantities are distance-independent and can be observed with spectroscopy and photometry respectively. From these, we can derive the effective radius which is then converted to a distance measurement.

The largest sample of FP distances to date is the one produced from the SDSS project (Howlett et al. 2022) containing about 34 000 distances. This sample recently surpassed the previous larger sample from the 6-degree Field Galaxy Survey peculiar velocity sample (6dFGSv, Springob et al. 2014) which contained nearly 9 000 distances. Figure 5.1 and 5.2 show the angular and redshift distribution of these two catalogues, respectively. The final estimated scatter of the SDSS sample is about 50% in distance moduli which corresponds to about 23% in distances, which is similar to the scatter of TF distances.

In section 5.3 I will present some of the growth-rate measurements using the 6dFGSv sample of distances. Currently there are no published results using the larger SDSS sample.

### 5.1.3 Type-Ia supernovae

As mentioned earlier, type-Ia supernovae (SNIa) have been used as distance indicators to measure the Hubble constant and the equation of state of dark energy. The state of the art results based on SNIa cover up to redshifts of 2.

Since SNIa provide distances, one can also estimate peculiar velocities by looking at the residuals with respect to the best-fit Hubble diagram. Peculiar velocities modifies mostly the observed redshifts than the observed magnitude. Magnitudes are only affected by the relativistic beaming, which scales luminosity distances as  $(1+z_{\text{pec}})^2$ , so it is a second-order effect. Figure 5.3 sketches how peculiar velocities would impact an ideal case of a SNIa Hubble Diagram without any kind of scatter. We can see how displacements are majoritarily along the

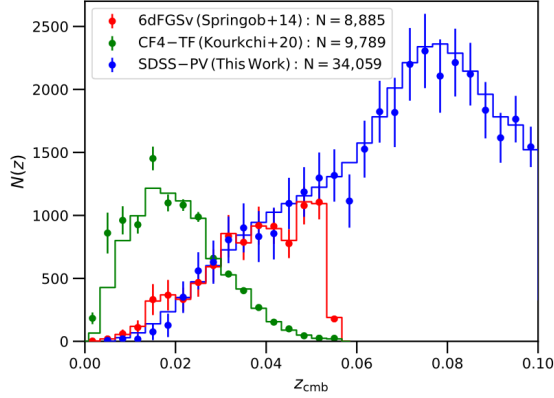


Figure 5.2: The redshift distribution of the distance samples: SDSS (blue) and 6dFGSv (red) from Fundamental Plane and Cosmicflows-4 (green) from Tully-Fisher relation. Points with error bars show the number of galaxies per bin of width 1000 km/s, where error bars include cosmic variance computed from mock catalogues. Solid lines are the mean redshift distributions from these mocks. Figure extracted from Howlett et al. 2022.

redshift direction.

Until now, peculiar velocities were mostly treated as nuisance when constructing the Hubble Diagram. Peterson et al. 2021 attempted to correct for peculiar velocities of some SNIa of the Pantheon+ sample (Dan Scolnic et al. 2022) by using reconstruction techniques on galaxy survey data and by identifying groups of galaxies. Peculiar velocities from SNIa have not been widely used to measure growth-rate for two main reasons. First, the small number of detected SNIa at low redshifts, where uncertainties in velocities are not prohibitive. Second, the inhomogeneity in the available samples which originate from several telescopes, cameras, each with its own observing strategy or prior selection of targets. As we will discuss later, selection effects are important when constructing velocity fields with any type of distance tracer.

The Zwicky Transient Facility (ZTF, Graham et al. 2019) is currently observing the whole available sky from the Northern hemisphere and will produce a large sample of about 5 000 SNIa at  $z < 0.1$  without any prior selection. This will be a homogeneous set ideal to perform statistical measurements with their peculiar velocities. Figure 5.4 compares the angular distribution of SNIa of the Pantheon+ sample to the ZTF Year 1 sample. Figure 5.5 compares the redshift distribution of the same samples. In section 5.5, I will present the plan of the project of measuring the clustering of ZTF peculiar velocities in combination with galaxies from DESI.

## 5.2 Biases affecting velocities

One of the main difficulties with distance measurements based on standardisable candles and their statistical properties (see next section) is the impact of certain types of biases. The review by Strauss and Willick 1995 discuss all these biases and how they impact clustering analyses. They can be classified in three types: selection bias, homogeneous and inhomogeneous Malmquist biases. The term Malmquist is often used to refer to selection bias, but I will follow the convention of Strauss and Willick 1995.

The selection bias, as the name indicates, is caused by some selection or detection cri-

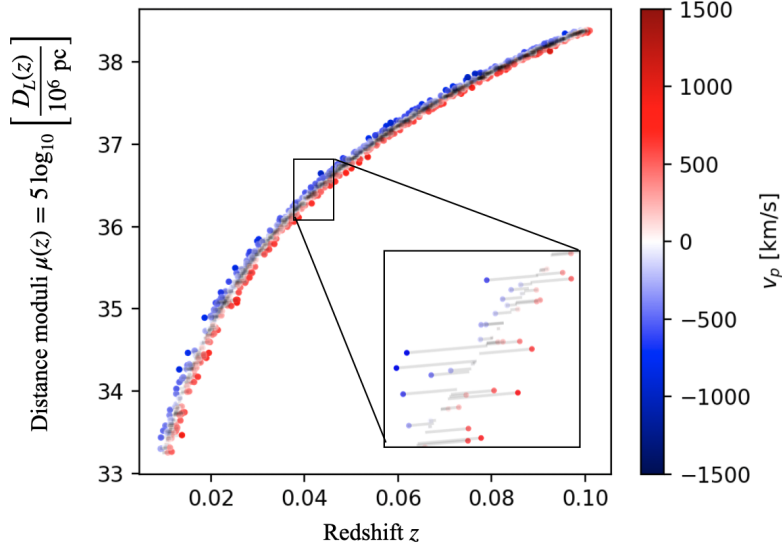


Figure 5.3: Effect of peculiar velocities on an ideal case of a Hubble Diagram of type-Ia supernovae without any intrinsic scatter or observational uncertainties. The displacements along the redshift direction are proportional to  $(1 + z_{\text{pec}})$  while those along the distance moduli direction are proportional to  $\log_{10}(1 + z_{\text{pec}})$ . The colour bar indicates the simulated peculiar velocity of each SNIa.

teria on the sample of objects. For instance, it is pretty common to impose some cuts on signal-to-noise ratio of fluxes or some magnitude limit. For objects close to the boundary, only the brighter objects pass the cuts. If one is interested in, say, the average distance of these objects, we will observe a bias relative to the true average distance. The intrinsic scatter in luminosity of all standardisable candles worsens this type of bias and it complicates any attempts to correct for these biases. Typical analyses of SNIa often use realistic simulations of the data, including the physical properties of SNIa that create the scatter and instrumental noise, to model these selection effects. The analysis of the Pantheon SNIa sample used a technique named Bayesian Estimation Applied to Multiple Species (BEAMS, Kunz et al. 2007) with Bias Corrections (BBC, Kessler and D. Scolnic 2017). This technique also accounts for mis-classified supernovae.

The homogeneous Malmquist bias manifests itself due to the fact that the number of objects observed as a function of distance to us increases with distance squared. When averaging objects in a given distance or redshift range, they are not uniformly distributed inside this range, so the average distance is higher than the true distance, often assigned to the center of the range. Is this true? Check!

The inhomogeneous Malmquist bias is a similar effect but generalised to the fluctuations in density (due to structures) along the line of sight. This bias depends on the precise shape of  $\delta(\hat{n}, d)$  where  $\hat{n}$  is a given direction in the sky and  $d$  is the distance to us.

It is possible to correctly account for all these biases in a fully Bayesian treatment of probability density functions of the distances (Graziani et al. 2019; Boruah et al. 2021).

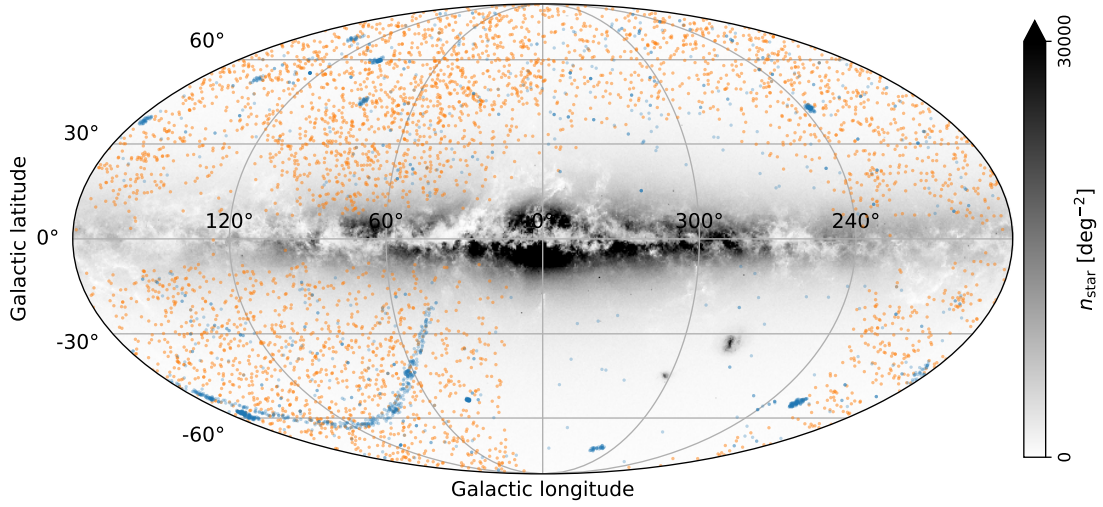


Figure 5.4: Angular distribution of type-Ia supernovae in Galactic coordinates for the Pantheon+ sample (blue) and for ZTF Year 2 sample (orange). On the background we see the stellar density distribution from Gaia.

### 5.3 Growth-rate measurements with velocities

With a set of distance or peculiar velocity measurements, we can study them statistically and compare to predictions from cosmological models, just as we did with the overdensities of fluxes in chapter 3 or galaxies in chapter 4.

One can see this additional set of observables as a scalar field, the radial velocity field  $v_r(\vec{x}_i)$ , sampled at the positions of the galaxies from which they were measured  $\vec{x}_i$ , where  $i = 1, \dots, N_{\text{gal}}$ . Instead of radial velocities, it is quite common to use the log of the ratio between the distance inferred from the redshift and the distance inferred from the luminosity, called log-distance ratio. Log-distance ratios have uncertainties that can be approximated by a Gaussian, whereas in velocity space the uncertainties become log-normal are harder to handle.

The first statistic that one can think of measuring is the correlation function of the radial velocities  $\xi_{v_r, v_r} \equiv \langle v_r(\vec{x}_i) v_r(\vec{x}_j) \rangle$  (or the equivalent for log-distance ratios). If one has, in addition to the velocities, a galaxy survey over the same volume, one can use the observed overdensity field  $\delta(\vec{x})$  to cross-correlate with the observed velocity field, i.e.,  $\xi_{\delta, v_r} \equiv \langle \delta(\vec{x}) v_r(\vec{y}) \rangle$ . The auto-correlation of the overdensity field  $\xi_{\delta, \delta} \equiv \langle \delta(\vec{x}) \delta(\vec{y}) \rangle$  is the equivalent to the one used in BAO and RSD measurements. These three statistics can be analysed jointly when comparing to models, increasing the constraining power.

I will shortly describe a few methods available in the literature that allows us to connect models to the observed densities and velocities, at the field level or at the two-point function level.

#### 5.3.1 Maximum likelihood method

The idea is to maximise the likelihood, assumed to be a multivariate Gaussian on the velocities,

$$\mathcal{L}(\vec{p}) = \frac{1}{(2\pi)^{\frac{n}{2}} \det[C(\vec{p})]^{1/2}} \exp \left[ -\frac{1}{2} \mathbf{v}_r^T C(\vec{p})^{-1} \mathbf{v}_r \right] \quad (5.2)$$

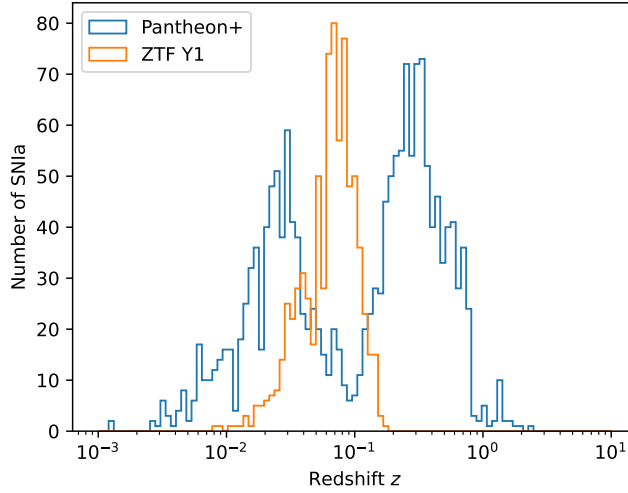


Figure 5.5: Redshift distribution of host galaxies of type-Ia supernovae for the Pantheon+ sample (blue) and for ZTF Year 1 sample (orange).

with respect to a set of parameters  $\vec{p}$ . Here  $v_r$  is a vector containing each measured peculiar velocity and the covariance matrix  $C(\vec{p})$  is the model for the two-point statistics of  $v_r$ . Each element of the covariance is given by  $C_{ij} = \xi_{v_r, v_r}^{\text{model}}(\vec{x}_i, \vec{x}_j, \vec{p}) + \delta_{ij}^K \sigma_i \sigma_j$ , where  $\sigma_i$  is a diagonal noise term, often accounting for instrumental uncertainties as well as intrinsic scatter of distance measurements. The first term in the covariance is the cosmic variance and it is connected to a cosmological model.

The cosmological part of the covariance  $C_{ij}$  is often written as a Fourier transform of the two-point statistics in Fourier space, often described by the velocity or the velocity-divergence power spectrum. Most of previous works have used linear perturbation theory with some empirical terms that account for non-linearities. In linear theory, the velocity field  $\vec{v}(\vec{x})$  is irrotational, meaning it is completely defined as the gradient of a scalar field  $\theta(\vec{x})$ . This velocity divergence is equal to the overdensity of matter  $\delta(\vec{x})$  for linear perturbations. Therefore the two-point statistics of the velocity is proportional to the linear matter power spectrum (and thus to  $\sigma_8^2$ ) and to the linear growth-rate of structures  $f$ .

The maximum likelihood method has been used on several peculiar velocity samples such as the 6dFGSv (Campbell et al. 2014; Springob et al. 2014) with velocities alone (Johnson et al. 2014), in combination with density field (Adams and Blake 2017; Adams and Blake 2020), or adding a few SNIa (Huterer et al. 2017).

At CPPM, Bastien Carreres is currently employing this technique to analyse the ZTF sample of type-Ia supernovae. I have implemented the basic algorithm and Bastien continued its development, adapting it to work with log-distance ratio instead of velocities, several mesh assignment schemes and models of clustering. He also implemented a whole pipeline to produce realistic simulations of ZTF SNIa light-curves, which can be painted on n-body simulations, and to analyse them as real data. He is currently studying the impact of selection biases on the estimates of  $f\sigma_8$  using the maximum likelihood method. First results on simulated data show that an unbiased measurement is possible if we limit our analysis to SNIa at  $z < 0.05$ . We will apply our methodology to real data once the final ZTF SNIa is set.



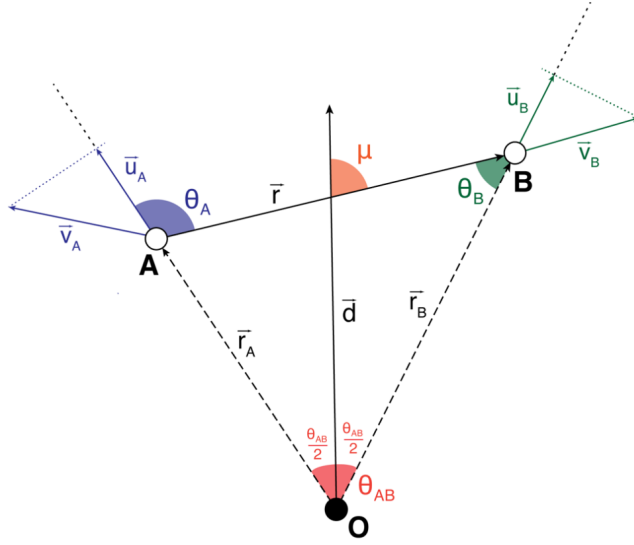


Figure 5.6: Non-parallel lines-of-sight of a pair of galaxies A and B as seen by an observer O.  $\vec{r}_A$  and  $\vec{r}_B$  are their position vectors, respectively, and  $\vec{r}$  is the separation vector between them. The peculiar velocities of galaxies A and B are represented by  $\vec{v}_A$  and  $\vec{v}_B$ , and the radial component of these velocities are represented by  $\vec{u}_A$  and  $\vec{u}_B$ . The pair line-of-sight  $\vec{d}$  is defined at the mid-angle and  $\mu \equiv \hat{d} \cdot \hat{r}$ . Figure extracted from Turner et al. 2023.

### 5.3.2 Compressed two-point statistics

The compression of two-point statistics means that we combine measurements into separation or wavenumber bins, depending if we are in configuration or Fourier space. Estimating compressed statistics is the most common method for standard clustering analysis, as we saw in earlier chapters, where we aim to estimate the correlation function or the power spectrum. The modelling happens at the compressed statistics level, so there are generally less degrees of freedom than for the maximum likelihood method.

In configuration space, the velocity auto correlation cannot simply be written as a unique function of separation  $\vec{r}$  between two galaxies, since velocities are a vector field. K. Gorski 1988 introduced the correlation tensor defined as

$$\Psi_{i,j}(\vec{r}_A, \vec{r}_B) \equiv \langle v_i(\vec{r}_A) v_j(\vec{r}_B) \rangle \quad (5.3)$$

where  $i, j$  denote each component of the peculiar velocity field  $\vec{v}$ . If the field is irrotational, homogeneous and isotropic, the correlation tensor can be written as

$$\Psi_{i,j}(\vec{r}) = [\Psi_{\parallel}(r) - \Psi_{\perp}(r)] \hat{r}_{Ai} \hat{r}_{Bj} + \Psi_{\perp}(r) \delta_{ij}^K \quad (5.4)$$

where  $\Psi_{\parallel}(r)$  and  $\Psi_{\perp}(r)$  are functions describing the correlations the components of the peculiar velocities parallel and perpendicular to the separation vector  $\vec{r}$ . Both functions can be written as transforms of the linear matter power spectrum.

The correlations between the radial velocities of two galaxies relative to the observer, as depicted in Figure 5.6, can be written as

$$\langle u_A(\vec{x}) u_B(\vec{x} + \vec{r}) \rangle = \Psi(r) \cos \theta_{AB} + [\Psi_{\parallel}(r) - \Psi_{\perp}(r)] \cos \theta_A \cos \theta_B. \quad (5.5)$$

K. M. Gorski et al. 1989 developed estimators that we can relate to a model that depends on  $\Psi_{\parallel}$  and  $\Psi_{\perp}$ . This type of correlation functions were used in Wang et al. 2018; Dupuy et al. 2019; Turner et al. 2021; Turner et al. 2023.

In Fourier space, given the sparsity of the data, it is hard to use FFTs on the observed velocity field. This is because voxels without data cannot be assigned zero velocity: the field would be ill-defined there. For this reason, Howlett 2019 suggests to use the momentum power spectrum, a quantity commonly used in the analysis of velocities in n-body simulations. The momentum is simply the mass-weighted velocity field, written as  $\vec{p}(\vec{x}) = [1 + \delta(\vec{x})] \vec{v}(\vec{x})$ . In the absence of galaxies or velocity measurements, the momentum is well defined and equal to zero.

The first application of the momentum power spectrum method was performed by Qin et al. 2019 on the combined sample of 6dFGSv and the 2MASS Tully-Fisher survey (2MTF).

The analysis in compressed statistics is generally faster and more intuitive than other methods. The data versus model comparison is commonly easier with compressed statistics. However, one has to be careful when accounting for survey properties, wide-angle effects or redshift evolution. These effects are generally easier to deal with in other methods.

### 5.3.3 Density-velocity comparison

The link between density and velocity fields is straightforward on quite large scales, where linear theory is valid. Another method to test for the consistency of the standard cosmological model is to compare the observed (radial) velocity field with predictions based on the observed density field. This method is often referred to as density-velocity comparison.

Reconstruction techniques discussed in section 4.2 aim to compute past trajectories of galaxies and therefore their velocities based on the observed density field. Reconstruction yields a prediction for the radial velocities at the locations where we measured them from the distance surveys. The inputs are often some value for the large scale bias  $b$  of the galaxy sample as well as the growth-rate  $f$ . Commonly, an empirical amplitude  $A$  is fitted to the relation  $v_r^{\text{obs}} = A v_r^{\text{model}}$ . Any deviations from  $A = 1$  would indicate an inconsistency in the standard cosmological model. From this amplitude, one can quote a value for  $f\sigma_8$  if a bias value and a  $\sigma_8$  value are assumed.

Several applications of this method can be found in the literature. Davis et al. 2011 used the Two Micron All-Sky Redshift Survey (2MRS or 2M++ for its extended version) galaxy catalogue as the input for reconstruction and fitted the inverse Tully-Fisher relation to spiral galaxies from the Spiral Field I-band ++ survey (SFI++). Carrick et al. 2015 improved the analysis and reported results using the same samples. Boruah et al. 2020 added peculiar velocities from 465 SNIa to the SFI++ and the 2MTF samples. Said et al. 2020 used a larger dataset of nearly 16k peculiar velocities by combining 6dFGSv and SDSS samples, and compared to the density field from the 2M++.

At CPPM, Elena Sarpa has a more advanced reconstruction technique called extended Fast Action Minimization (eFAM, Sarpa et al. 2021) which can be readily used to implement the density-velocity comparison. The eFAM technique yields higher-order trajectories than Zeldovich reconstruction, which might lead to more realistic predictions for the peculiar velocities. This technique will be potentially applied to the DESI and ZTF data, after a detailed comparison with the maximum likelihood method.



### 5.3.4 Forward-model of density and velocity field

Instead of estimating summary statistics with the density and velocity fields, one can attempt to fit directly the observed fields with some model that depend on cosmological parameters. Typically, the model is composed of some cosmological parameters and a set of initial density and velocity fields (one amplitude parameter per mode). These initial conditions are then evolved linearly or using more advanced techniques and then compared to observations. The final product is a posterior on cosmological parameters and on the amplitudes of the initial fields. This method is commonly referred to as *forward-modelling*.

In forward modelling, complications due to selection functions are slightly easier to deal with in principle. One has to simulate observations on the evolved fields just as one does when creating mock catalogues. Also, by modelling the observed field entirely, we capture more information from higher-order statistics than just the two-point function. The inconvenience of this method is the computing time required to perform the inference, given the large number of free parameters and the time to evolve each sample from initial conditions.

Lavaux 2016 implemented the forward-modelling applied to both density and velocity fields. A similar approach was used on CosmicFlows3 data by Graziani et al. 2019. In Boruah et al. 2021 they discuss the treatment of the inhomogeneous Malmquist bias in the Bayesian formalism (see section 5.2). In Prideaux-Ghee et al. 2023, they used a Bayesian hierarchical modelling approach to fit both density and velocity fields using a Hamiltonian Monte Carlo sampling of the likelihood.

## 5.4 Current measurements

Several measurements were made in the past decade using a variety of distance surveys, galaxy surveys and methods. Table 5.1 summarises most measurements in the literature. They are also visually represented in Figure 5.7. All these measurements of the growth-rate  $f\sigma_8$  correspond to an effective redshift between zero and 0.1. These results are in good agreement with the predictions from a  $\Lambda$ CDM+GR model, with cosmological parameters from Planck Collaboration et al. 2020. However, one can notice how they do not have the same uncertainties, even if, for some cases, they use the same datasets. Few analyses quote systematic uncertainties which can be quite significant for some methods.

The goal of the project I am currently leading at CPPM is to deeply understand those variations while optimising the analysis. Also we plan to work closely to the data reduction teams in order to produce the best set of simulated data. Mock catalogues will be essential to estimate biases in best-fit parameters, statistical and systematic uncertainties. In the following section I will introduce two datasets that will be used in our project: DESI and ZTF.

## 5.5 DESI and ZTF

Both the Dark Energy Spectroscopic Instrument (DESI) and the Zwicky Transient Facility (ZTF) are currently observing similar area of the sky and building one of the best samples for a joint study of densities and peculiar velocities for cosmology. One of the main advantages is the great overlap in volume of both surveys. The 14k deg<sup>2</sup> footprint of DESI will be completely covered by ZTF observations, which cover up to 18k deg<sup>2</sup> of area.

The main science goal of DESI is to produce precise measurements of the Universe's expansion and the growth history of structures, in order to test models of dark energy and gravity. DESI will observe more than 20 million galaxies and quasars from up to redshifts of 4.5.

Table 5.1: Measurements of the growth-rate of structures from peculiar velocity and galaxy survey data.

Reference	Dataset	Method	$f\sigma_8$
Johnson et al. 2014	6dFGSv	Max-likelihood	$0.428^{+0.079}_{-0.068}$
Huterer et al. 2017	6dFGSv	Max-likelihood	$0.428^{+0.048}_{-0.045}$
Adams and Blake 2020	6dFGSz,6dFGSv	Max-likelihood	$0.384 \pm 0.052^{(\text{stat})} \pm 0.061^{(\text{syst})}$
Lai et al. 2023	SDSS-FP	Max-likelihood	$0.405^{+0.076(\text{stat})}_{-0.071} \pm 0.009^{(\text{syst})}$
Qin et al. 2019	6dFGSv,2MTF	Compressed 2pt	$0.404^{+0.082}_{-0.081}$
Howlett et al. 2017b	2MTF	Compressed 2pt	$0.510^{+0.090}_{-0.080}$
Nusser 2017	CF3,2MRS	Compressed 2pt	$0.400 \pm 0.080$
Dupuy et al. 2019	CF3	Compressed 2pt	$0.430 \pm 0.030^{\text{obs}} \pm 0.110^{\text{cosmic}}$
Davis et al. 2011	2MRS,SFI++	$\delta - v$ comparison	$0.310 \pm 0.040$
Carrick et al. 2015	2MRS,SFI++	$\delta - v$ comparison	$0.401 \pm 0.024$
Boruah et al. 2020	2M++,2MTF,SFI++,A2	$\delta - v$ comparison	$0.400 \pm 0.017$
Said et al. 2020	2M++,6dFGSv,SDSS-FP	$\delta - v$ comparison	$0.311 \pm 0.027$

The bright time of the survey, i.e., when the moon is up, is dedicated to building a flux limited sample of galaxies at low redshift: the Bright Galaxy Survey (BGS). The BGS will collect roughly 10 million galaxies with  $r$ -band magnitude  $r < 19.5$ , which lie at  $z < 0.6$ . A fainter  $19.5 < r < 20.175$  sample will extend and increase the completeness in stellar mass of the BGS sample. This sample is not only ideal for BAO and standard RSD measurements, but also for higher order statistical measurements, multi-tracer analyses, given its high completeness and simple selection function. Most of the ZTF SNIa will have their host-galaxy redshifts measured by the DESI BGS. It is expected that the BGS observations will complete by 2024 the latest.

The ZTF survey main goal is to observe the transient sky in three optical bands. Most of the sky area is covered every two nights with a given filter, which is excellent cadence for discovery and study of SNIa. At its current configuration, ZTF is expected to discover around 15k supernovae, out of which about 30% will have spectroscopic classifications. The host galaxy redshifts are obtained from archival data, some of which will be observed by DESI in the near future. We expect a final sample of 5k cosmology-level spectroscopic SNIa at  $z < 0.1$ .

Figure 5.8 compares the density of tracers between DESI BGS and ZTF SNIa, expected at the end of their 5-year programs. The density of ZTF SNIa falls quickly after  $z \sim 0.06$  where selection effects start to kick in. Rough estimates show that the ZTF SNIa sample is *complete*, i.e., all potentially observable SNIa were observed, at  $z < 0.05$ . While the density of ZTF SNIa is several orders of magnitudes lower than those of BGS galaxies, the fact that each SNIa yields a measurement of the velocity field makes them a powerful tracer nevertheless.

I studied the constraining power of DESI BGS and ZTF SNIa using the Fisher formalism. I computed the uncertainty expected on  $f\sigma_8$  from a RSD analysis with DESI BGS galaxies by themselves or from a joint RSD + PV analysis of DESI BGS and ZTF SNIa. I employed the same Fisher forecast code<sup>1</sup> described in Howlett et al. 2017a. The inputs of the calculation are the expected comoving density of tracers for galaxies and peculiar velocities at the end of both surveys, the intrinsic scatter of the distance indicators parametrised by  $\alpha \equiv \sigma(D_L)/D_L \sim 10$  or 20%, and the scale range (in Fourier space) assumed for the measurement parametrised by  $k_{\text{max}} \sim 0.1$  or  $0.2 \text{ hMpc}^{-1}$ . The model for the galaxy clustering is a simple linear RSD model with galaxy bias  $b$  and a Gaussian Fingers-of-God term function of a galaxy velocity

<sup>1</sup>[https://github.com/CullanHowlett/PV\\_fisher](https://github.com/CullanHowlett/PV_fisher)

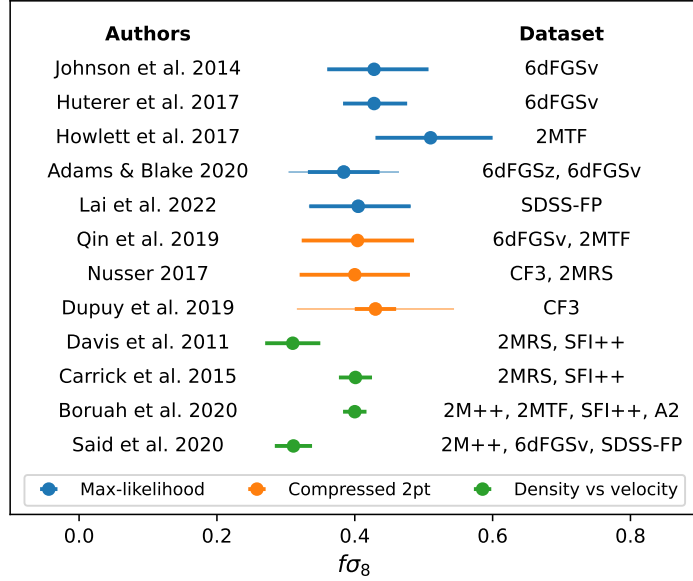


Figure 5.7: Measurements of the growth-rate of structures  $f\sigma_8$  from peculiar velocity and galaxy survey data. The references and values are listed in Table 5.1. The light error bars include systematic errors. In the case of Dupuy et al. 2019, the extra contribution is from cosmic variance, instead.

dispersion  $\sigma_v$ . The model for the velocity correlations is similar, scaling with  $f$ , but with a different empirical correction for the fact we observe velocities in redshift-space (Koda et al. 2014). This empirical function depends on an extra dispersion parameter  $\sigma_u$ . A total of four free parameters define the model. I focused on the constraints on  $f\sigma_8$  by using tracers over  $0 < z < 0.1$ , which corresponds to a measurement at an effective redshift  $z_{\text{eff}} = 0.08$ .

Results of the Fisher forecast are shown in Table 5.2 for different samples and analysis choices. First, we can see how the uncertainties on  $f\sigma_8$  are significantly reduced when including ZTF SNIa, particularly when the intrinsic scatter is small and when using a larger range of scales. We see an improvement of a factor of 2.3 at most. Second, we see how increasing the scale range  $k_{\text{max}}$  also improves constraints. This is why it is important to be able to model the clustering of galaxies and peculiar velocities on non-linear scales. Third, we see how improving the standardisation of SNIa also helps reducing uncertainties on  $f\sigma_8$ . In conclusion, a simple Fisher forecast predicts a 9% measurement of  $f\sigma_8$  in the optimistic scenario and a 20% in the pessimistic scenario.

## 5.6 Growth-rate measurement from simulated ZTF data

Bastien Carreres, PhD student I co-supervise with Benjamin Racine and Dominique Fouchez at CPPM, is currently developing the whole analysis chain to measure the growth-rate of structures from peculiar velocities derived from ZTF type-Ia supernovae (SNIa) data. He has been developing a code based on the maximum-likelihood method (section 5.3.1) and validating it with the most realistic set of simulated ZTF data that he also produced. In this section, I summarise his findings which will be detailed in Carreres et al. (in prep).

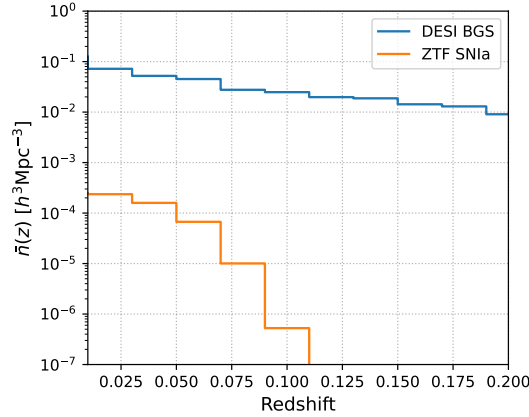


Figure 5.8: Comoving density of tracers as a function of redshift for the galaxies of the DESI Bright Galaxy Survey (blue) and for type-Ia supernovae from ZTF (orange). Densities were computed assuming a flat  $\Lambda$ CDM model with Planck 2018 best-fit parameters, based on simulated data and scaled to match numbers expected at the end of both programs. Both samples overlap in redshift and in sky coverage, which is ideal for joint studies of densities and peculiar velocities.

### 5.6.1 Simulating ZTF type-Ia supernovae with peculiar velocities

The goal of the simulations is to create a set of light-curves from SNIa which contain realistic and correlated peculiar velocities.

The best source of peculiar velocities are n-body simulations. We currently use two sets of large-scale cosmological n-body simulations with matter only. The first set is named Dark Energy and Massive Neutrino Universe (DEMNUi, Castorina et al. 2015) aimed at account correctly for the impact of massive neutrinos in the formation of structures. They were produced with the tree particle mesh-smoothed particle hydrodynamics (TreePM-SPH) code Gadget-3 with a recipe given by Viel et al. 2010 to simulate massive neutrinos. The DEMNUi boxes are  $2 h^{-1} \text{Gpc}$  on a side, and contain  $2048^3$  matter particles and  $2048^3$  neutrino particles. Initial conditions are set at  $z = 99$  and evolved assuming a flat  $\Lambda$ CDM universe with parameters from Planck 2013 results (Collaboration et al. 2014). Different simulations were ran with varying mass of neutrino species (0, 0.16, 0.32, and 0.48 eV) and varying equations-of-state for dark energy. Haloes were defined using a Friend-of-friends algorithm, assembling a minimum number of 32 matter particles. This first set of simulations is not yet public and access it only granted by demand to the authors. The second set of simulations we use are the AbacusSummit suite (Maksimova et al. 2021), which use the Abacus code (Garrison et al. 2021) to produce 150 boxes of  $2 h^{-1} \text{Gpc}$  on a side, containing  $6912^3$  particles and spanning 97 cosmological models. These cosmological models have different values for  $\omega_b$ ,  $\omega_{\text{cdm}}$ ,  $h$ ,  $A_s$ ,  $n_s$ ,  $\alpha_s$ ,  $N_{\text{eff}}$ ,  $w_0$  and  $w_a$ . Initial conditions are created using first order Lagrangian perturbation theory with the public code ZELDOVICH-PLT<sup>2</sup>. Halos are found using the CompaSO technique (Hadzhiyska et al. 2022). These simulations and several data products are publicly available and documented<sup>3</sup>. The AbacusSummit suite is currently the state-of-the-art in the market and are currently being used by the DESI collaboration.

<sup>2</sup><https://github.com/abacusorg/zeldovich-PLT>

<sup>3</sup><https://abacussummit.readthedocs.io/>

Table 5.2: Fisher forecast on measurements of the growth-rate of structures times the normalisation of the power spectrum  $f\sigma_8(z_{\text{eff}}) = 0.45$  where  $z_{\text{eff}} = 0.08$  for the complete samples of DESI BGS and ZTF type-Ia supernovae. Results are shown for different analysis choices:  $k_{\text{max}}$  is the maximum wavenumber used in the analysis,  $\sigma(D_L)/D_L$  is the fractional error on luminosity distances due to intrinsic scatter. The assumed number densities of tracers is the one shown in Figure 5.8. The total assumed sky area with overlapping coverage from both experiments is 14k deg<sup>2</sup>.

Dataset	$k_{\text{max}}$	$\sigma(D_L)/D_L$	$\sigma(f\sigma_8(z_{\text{eff}}))/f\sigma_8(z_{\text{eff}})$
DESI BGS	0.1	-	0.58
DESI BGS	0.2	-	0.21
ZTF SNIa	0.1	0.05	0.22
ZTF SNIa	0.1	0.10	0.35
ZTF SNIa	0.2	0.05	0.19
ZTF SNIa	0.2	0.10	0.32
DESI BGS + ZTF SNIa	0.1	0.05	0.12
DESI BGS + ZTF SNIa	0.1	0.10	0.20
DESI BGS + ZTF SNIa	0.2	0.05	0.09
DESI BGS + ZTF SNIa	0.2	0.10	0.13

ZTF observes spectroscopically classified SNIa up to redshifts of 0.1, which corresponds to a comoving distance of  $\sim 293 h^{-1}\text{Mpc}$  in a Planck 2018 best-fit cosmology. This means that we can subdivide a  $(2 h^{-1}\text{Gpc})^3$  simulation box into roughly 9 ZTF SNIa volumes (full-sky, even though ZTF observes the Northern sky). Because of large-scale correlations, these sub-boxes are not completely independent of each other. We place the observer at the center of each sub-box and compute the radial component of peculiar velocities for each halo. Eventually we convert comoving distances into redshifts assuming a fiducial cosmology. At this step, we can also include the effect of redshift-space distortions (RSD) using the radial velocities. Currently we do not apply any halo mass cuts in our halo samples.

With halo samples in hand, we randomly assign SNIa events following a given rate of events per volume, e.g.,  $(2.35 \pm 0.24) \times 10^4 \text{Gpc}^{-3} \text{yr}^{-1}$  (Perley et al. 2020). Using observation logs from real ZTF observations, we can know which exposures could observe our SNIa events within some time window comparable to the duration of a SNIa light-curve. For each observed SNIa, we draw stretch and colour parameters defined as in the SALT2 formalism (Guy et al. 2007) and following realistic distributions (D. Scolnic and Kessler 2016; Nicolas et al. 2021). Intrinsic scatter is added by hand on the peak magnitudes. The fluxes are simulated using the SNCOSMO package<sup>4</sup>. Each exposure contains information about the seeing, airmass, and noise, which are then combined to obtain fluxes and their uncertainties. The peculiar velocity information is included in the observed redshift of the SNIa event and accounted for when deriving fluxes in each band.

Once light-curves are computed for all observed SNIa, we simulate selection effects. First, we define a photometric cut to emulate detection. We exclude any light-curve without two epochs with flux signal-to-noise ratio above 5. Second, we emulate the spectroscopic follow-up for classification. The selection for spectro follow-up is essentially a magnitude cut (roughly  $r < 18.5$ ), which ensures enough signal in spectra. This second cut is the main responsible for the limit of  $z \sim 0.1$  of the SNIa sample as seen in Figure 5.8. A mock ZTF SNIa sample

<sup>4</sup><https://sncosmo.readthedocs.io/>

with peculiar velocities is the result of this procedure. The package producing these mock catalogues is publicly available and documented<sup>5</sup>.

### 5.6.2 Measuring peculiar velocities and the growth-rate

With a realistic mock catalogue of ZTF SNIa data in hand, we proceed to the measurement of their peculiar velocities and the estimate of the growth-rate of structures  $f\sigma_8$ .

Figure 5.9 shows a few examples of light-curves from mock ZTF data. These light-curves are then fitted with the same SALT2 models used to produce them. The fits yield values for the stretch, colour, peak brightness and the time of peak brightness, with their uncertainties. A Hubble diagram is fitted including terms accounting for correlations between brightness and stretch/colour (Tripp relation). Since we are at such low redshifts, we can model  $H(z)$  as a linear function since the cosmological dependency with  $\Omega_m$  or  $\Omega_\Lambda$  only kicks in at  $z \sim 0.4$ , much higher than our sample. The peculiar velocity estimator uses the residuals of the observed distance moduli to the best-fit Hubble diagram. While most of the displacements due to velocities happen in the redshift direction, the noise is much larger in the distance moduli direction, so the most convenient observable is the Hubble residual  $\Delta\mu \equiv \mu_{\text{obs}} - \mu_{\text{model}}(z_{\text{obs}})$ . Using a linear expansion of the Hubble diagram, one can write an approximation of the corresponding peculiar velocity, though its noise properties become non-Gaussian. Therefore most studies simply use  $\Delta\mu$  as the peculiar velocity observable.

Figure 5.10 displays the residuals of the Hubble diagram fitted on simulated samples of ZTF light-curves, relative to the true input Hubble diagram. The average of eight independent realisations is also displayed with error bars. A clear bias relative to the input model is seen at  $z > 0.06$ , due mainly to the selection for spectroscopic follow-up. If the simulations are realistic enough, this bias is an indication that ZTF SNIa samples are *complete* up to  $z \sim 0.06$ , defining a so called *volume limited* sample. Above this limit, we need to correctly account for biases if one is interested in constraining cosmology with the Hubble diagram. In this work, we studied the impact of this selection bias in the measurement of peculiar velocities and subsequently on the estimates of the growth-rate  $f\sigma_8$ .

For our measurement of  $f\sigma_8$  using peculiar velocities, we are currently employing the maximum-likelihood method (section 5.3.1). Instead of using a linear matter power spectrum as a model for the velocity divergence power spectrum, we are trying to use improved models including regularised perturbation theory (RegPT, Taruya et al. 2012) and empirical formulas derived from n-body simulations (Bel et al. 2019). These models include some level of non-linearities which would allow us to use a larger range of scales when constraining growth. Also, linear theory is known to break at quite large scales  $k \sim 0.1$ , particularly at  $z \sim 0$ .

When building the model covariance matrix used in the likelihood (Eq. 5.2) for  $N_v$  measurements of peculiar velocities, the resulting matrix will have  $N_v^2$  elements, which can become prohibitively large and slow to compute. We expect that ZTF will produce a set of about 5000 SNIa which is still feasible, but larger sets require a different solution. One of the solutions is to assign the velocities to a mesh, effectively reducing the number of measurements. This also has the advantage of smoothing out a fraction of the non-linear clustering. The mesh smoothing is taken into account in the model covariance matrix.

There are three free parameters in the model covariance matrix:  $f\sigma_8$  which is simply to overall amplitude of velocity-velocity correlations,  $\sigma_v$  accounting for the remaining non-linear velocities,  $\sigma_u$  accounting for the change in the signal due to redshift-space distortions.

<sup>5</sup><https://snsim.readthedocs.io/>

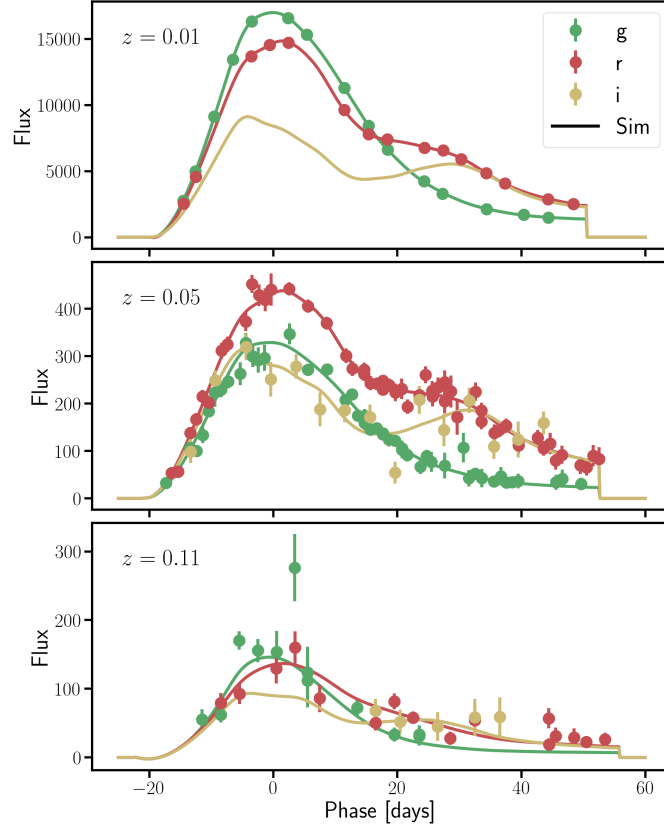


Figure 5.9: Three simulated type-Ia supernovae light-curves as observed by ZTF, ranging from low to mid and high redshift. The solid line represents the input light-curve model while the points with error bars are simulated observed fluxes including instrumental noise. The flux is in arbitrary units. Figure extracted from Carreres et al. (in prep).

We use the `IMINUIT` package<sup>6</sup> to find the maximum of the likelihood and to estimate uncertainties. We also explore the full posterior distribution using a Monte Carlo Markov Chain (MCMC).

Figure 5.11 shows one example of the posterior on the three fitter parameters. In order to avoid selection biases, we restrain our dataset to SNIa at  $z < 0.06$ , which yields about 2000 objects for this case. Contours show the 68 and 95% confidence levels. Green contours are the result of analysing the peculiar velocities directly from the halo catalogue, without any source of noise. Red contours are the result when including all the effects of a realistic dataset and analysis. Both measurements are consistent with the expected value for  $f\sigma_8$ , even though uncertainties are quite large.

Figure 5.12 shows how best-fit values of  $f\sigma_8$  change when increasing the dataset up to  $z_{\max}$ . We clearly see the impact of selection biases when considering SNIa beyond  $z \sim 0.08$ , based on the average of 8 mock realisations. This would indicate that an unbiased measurement of  $f\sigma_8$  is possible with ZTF SNIa if one restricts the analysis to the volume limited sample. Another interesting result is that the uncertainty in  $f\sigma_8$  does not seem to decrease significantly when increasing  $z_{\max}$ . This is simply a manifestation of the quick fall in the comoving density of SNIa beyond  $z \sim 0.08$ , as seen in Figure 5.8. Also in the same redshift range, a large

<sup>6</sup><https://iminuit.readthedocs.io/>

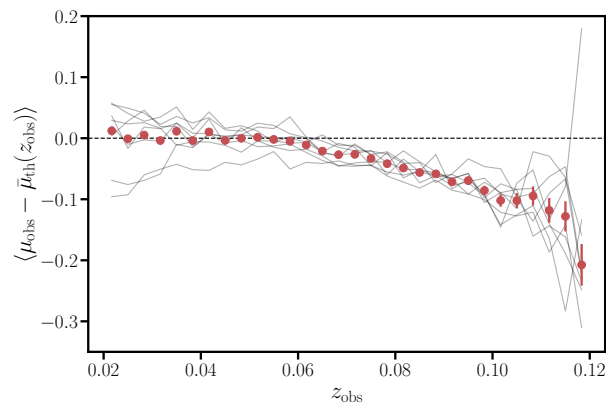


Figure 5.10: Distance moduli inferred from simulated ZTF SNIa light-curves compared to the input cosmological model. Grey lines show 8 different realisations while points with error bars are their average. Selection effects cause a bias in distance moduli estimates at  $z > 0.06$ . Figure extracted from Carreres et al. (in prep).

set of DESI galaxies will be available to perform a joint density-velocity analysis.

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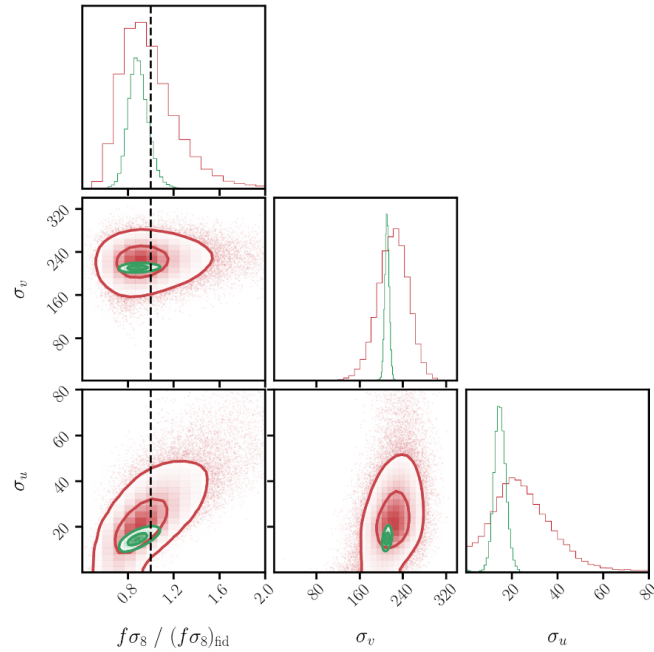


Figure 5.11: Example of posterior probability distribution for the growth-rate  $f\sigma_8$  and the dispersion parameters  $\sigma_v$  and  $\sigma_u$ , obtained from an analysis on a mock ZTF sample. Green contours are the result when halo radial peculiar velocities are used directly without any source of noise. Red contours are the case of a full ZTF mock SNIa catalogue, including light-curve simulation and fitting. The dashed vertical line represents the input value of the simulation. Figure extracted from Carreres et al. (in prep).

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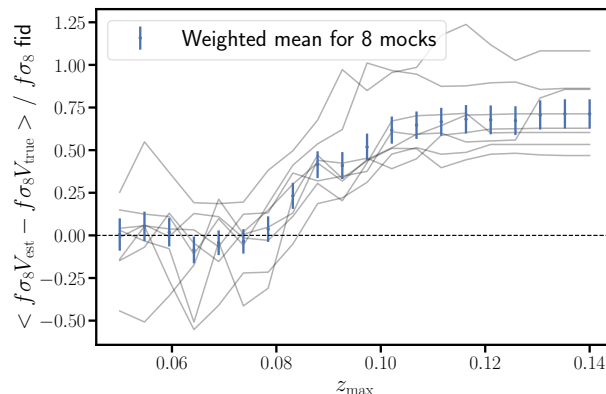


Figure 5.12: Estimates of the growth-rate  $f\sigma_8$  from mock ZTF SNIa data. We compare the measurements between the estimated peculiar velocities to the same measurements performed on the true velocities. Each measurement considers SNIa up to maximum redshift  $z_{\text{max}}$ . The selection bias from Figure 5.10 manifests itself on growth-rate estimates from  $z_{\text{max}} > 0.08$ . Figure extracted from Carreres et al. (in prep).

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# Conclusion and future outlook

The last 11 years of my life were dedicated to the quest of dark energy using data from the best spectroscopic galaxy surveys available, such as the Sloan Digital Sky Survey Baryon Oscillation Spectroscopic Survey (BOSS), its extended version (eBOSS) and the Dark Energy Spectroscopic Survey (DESI). I have been lucky to experiment with data from a diverse range of redshifts. In the last years of this exploration, I had the privilege to work with excellent people at different levels: undergrad (Andrei Marin), masters (Nattapon Preedasak), PhD candidates (Samantha Youles, Bastien Carreres, Vincenzo Aronica, Tyann Dumerchat) and postdocs (Elena Sarpa), as their main or co-supervisor. Many others were also part of it for shorter durations but equally enjoyable.

At high-redshift, I used quasars and their  $Ly\alpha$  forests to measure baryon acoustic oscillations (BAO). This measurement is still the highest redshift one to date, if we exclude the cosmic microwave background. Thanks to my experience with analysis spectra, I got involved in the details of the eBOSS data reduction pipeline, where my developments ultimately improved cosmological constraints. With S. Youles, we explored a new type of systematic effect in the correlation function of quasars and forests, which should be further studied in future DESI data analyses.

At mid-redshifts, I used mainly luminous red galaxies to measure their clustering and derive constraints from BAO and redshift-space distortions (RSD). Our results with eBOSS data are still the current state-of-the-art in our field, and will continue to be until we publish the first results with the DESI Year 1 sample. With T. Dumerchat and V. Aronica, we explored a joint analysis of Fourier and configuration space clustering for BAO and RSD. V. Aronica is also currently leading the effort of understanding the impact of photometric systematics on DESI data. I explored 21cm radio data from the Green Bank Telescope, and worked on its cross-correlation with eBOSS galaxies, which will be useful for future data from surveys such as the Square Kilometre Array.

At low-redshifts, peculiar velocities can be directly measured using type-Ia supernovae. I turned my interest in the past two years to the joint analysis of galaxy clustering and peculiar velocities, in order to measure the growth-rate of structures. With B. Carreres, we are exploring methods to measure growth from peculiar velocity data from SNIa. He produced the most realistic sets of mock catalogues for the Zwicky Transient Facility. T. Dumerchat is studying emulators for a better modelling of the two-point functions in the non-linear regime, while E. Sarpa is testing the density-velocity method based on reconstruction techniques.

During those past years, I gathered experiences with a diverse range of cosmic epochs, cosmic probes, and analysis techniques. I had the opportunity of collaborating with a large team of researchers and more recently I supervised the work of PhD students and postdocs. This experience will be essential to face the challenges of the next decades, when experiments will push the precision of measurements to its limit.

## Future outlook

The future in cosmology is very promising and I hope to continue providing key contributions to the field, as a Professor of Aix-Marseille University and the *Centre de Physique des Particules de Marseille*.

In the next four years, DESI will continue its observing program and will be the first stage-IV experiment to be complete, producing the best three-dimensional map of the Universe over  $0 < z < 4$ . The planning of its follow-up survey, DESI-II, has already started, and our team could be involved in its development and data analysis. The instrument would be the same but our target a new set of targets for scientific goals that complement those from DESI. The MegaMapper is a stage-V proposed experiment that would be based on a 6.5-metre telescope and equipped with more than 25 000 optical fibres in the focal plane. If approved, this project would probe the Universe over  $2 < z < 5$  to learn about dark energy and inflation.

The Euclid satellite is under construction and it is planned to be launched in 2023. Euclid is both a photometric and a spectroscopic survey, so it will produce weak-lensing and clustering measurements. The spectroscopic component will cover the range  $1 < z < 2$ , with a large overlap with the DESI sample. The combination of Euclid and DESI data will yield the best constraints on dark energy by the end of 2030. Lessons learned from DESI could be transferred to the analysis of Euclid data, and our team could be a part of this work.

The Rubin Observatory Legacy Survey of Space and Time (Rubin-LSST) will start commissioning and science observations by 2024. Thanks to its cadence and depth, Rubin-LSST will observe light-curves for  $10^5$  supernovae over  $0.1 < z < 0.4$ , and produce lensing-quality images over most of the available Southern sky. With our expertise on the analysis of ZTF type-Ia supernovae, we will be well positioned to analyse the larger sample of type-Ia supernovae provided by Rubin-LSST. The observational program will last at least ten years and the complementarity with spectroscopic surveys is a rich source of interesting projects we could lead.

I wish to continue working towards the goal of understanding dark energy, while forming new generations of scientists, in the role of Professor of Aix-Marseille University. I hope this manuscript convince readers of my capabilities as a researcher and as a supervisor of a research team.