



HAL
open science

Entropy-calibrated stellar modeling: Testing and improving the use of prescriptions for the entropy of adiabatic convection

L. Manchon, M. Deal, M.-J. Goupil, A. Serenelli, Y. Lebreton, J. Klevas, A. Kučinskas, H.-G. Ludwig, J. Montalbán, L. Gizon

► To cite this version:

L. Manchon, M. Deal, M.-J. Goupil, A. Serenelli, Y. Lebreton, et al.. Entropy-calibrated stellar modeling: Testing and improving the use of prescriptions for the entropy of adiabatic convection. *Astronomy and Astrophysics - A&A*, 2024, 687, pp.A146. 10.1051/0004-6361/202347700. hal-04639444

HAL Id: hal-04639444

<https://hal.science/hal-04639444>

Submitted on 10 Jul 2024










HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.



Distributed under a Creative Commons Attribution 4.0 International License

Entropy-calibrated stellar modeling: Testing and improving the use of prescriptions for the entropy of adiabatic convection[★]

L. Manchon^{1,2}, M. Deal^{3,4}, M.-J. Goupil², A. Serenelli^{5,6,7}, Y. Lebreton^{2,8}, J. Klevas⁹, A. Kučinskas⁹,
H.-G. Ludwig¹⁰, J. Montalbán^{11,12}, and L. Gizon^{1,13}

¹ Max-Planck-Institut für Sonnensystemforschung, Justus-von-Liebig-Weg 3, 37077 Göttingen, Germany

² LESIA, Observatoire de Paris, Université PSL, CNRS, Sorbonne Université, Université de Paris, 92195 Meudon, France
e-mail: louis.manchon@obspm.fr

³ LUPM, Université de Montpellier, CNRS, place Eugène Bataillon, 34095 Montpellier, France

⁴ Centro de Astrofísica da Universidade do Porto, Rua das Estrelas, 4150-762 Porto, Portugal

⁵ Institute of Space Sciences (ICE, CSIC), Carrer de Can Magrans S/N, 08193 Cerdanyola del Valles, Spain

⁶ Institut d'Estudis Espacials de Catalunya (IEEC), Carrer Gran Capita 2, 08034 Barcelona, Spain

⁷ Max-Planck-Institut für Astronomie, Königstuhl 17, 69117 Heidelberg, Germany

⁸ Université de Rennes, CNRS, IPR (Institut de Physique de Rennes), UMR 6251, 35000 Rennes, France

⁹ Institute of Theoretical Physics and Astronomy, Vilnius University, Saulėtekio al. 3, 10257 Vilnius, Lithuania

¹⁰ Zentrum für Astronomie der Universität Heidelberg, Landessternwarte, Königstuhl 12, 69117 Heidelberg, Germany

¹¹ Dipartimento di Fisica e Astronomia Augusto Righi, Università degli Studi di Bologna, Via Gobetti 93/2, 40129 Bologna, Italy

¹² School of Physics and Astronomy, University of Birmingham, Birmingham B15 2TT, UK

¹³ Institut für Astrophysik, Georg-August-Universität Göttingen, Friedrich-Hund-Platz 1, 37077 Göttingen, Germany

Received 10 August 2023 / Accepted 24 February 2024

ABSTRACT

Context. Modeling the convection process is a long-standing problem in stellar physics. To date, all ad hoc models have relied on a free parameter, α , (among others) that has no real physical justification and is therefore poorly constrained. However, a link exists between this free parameter and the entropy of the stellar adiabat. There are existing prescriptions, derived from 3D stellar atmospheric models, that treat entropy as a function of stellar atmospheric parameters (effective temperature, surface gravity, and chemical composition). This can offer sufficient constraints on α through the development of entropy-calibrated models. However, several questions have arisen as these models are increasingly used with respect to which prescription should be used and whether it ought to be used in its original form, along with the impacts of uncertainties on entropy-calibrated models.

Aims. We aim to study the three existing prescriptions in detail and determine which of them demonstrate the most optimal performance and how it should be applied.

Methods. We implemented the entropy-calibration method into the stellar evolution code (Cesam2k20) and performed comparisons with the Sun and the α Cen system. In addition, we used data from the CIFIST grid of 3D atmosphere models to evaluate the accuracy of the prescriptions.

Results. Of the three entropy prescriptions currently available, we determined the one that has the best functional form for reproducing the entropies of the 3D models. However, the coefficients involved in this formulation must not be taken from the original paper because they were calibrated against a flawed set of entropies. We also demonstrate that the entropy obtained from this prescription should be corrected for the evolving chemical composition and for an entropy offset different between various EoS tables. This must be done following a precise procedure to ensure that the classical parameters obtained from the models are not strongly biased. Finally, we provide a data table with entropy of the adiabat of the CIFIST grid, along with the fits for these entropies.

Conclusions. Thanks to a precise examination of entropy-calibrated modeling, we are able to offer our recommendations with respect to which adiabatic entropy prescription to use, how to correct it, and how to implement the method into a stellar evolution code.

Key words. convection – methods: numerical – stars: evolution – stars: fundamental parameters – stars: interiors

1. Introduction

Convection modeling is a major problem in the context of obtaining realistic stellar evolution models. Stellar convective regions are highly turbulent media where the Reynolds number can reach values on the order of 10^8 (solar photosphere; Komm et al. 1991) to 10^{14} (bottom of the solar convective envelope; Kupka & Muthsam 2017). In such regimes, Navier-Stokes

equations can only be solved numerically. Conducting such hydrodynamic simulations would require high spatial resolutions on a large domain going from the top of adiabatic region to the outer layers of the photosphere and over a timescale of stellar evolution. This is completely unachievable with the computational capabilities available presently and even in the foreseeable future.

To circumvent this problem, physicists have designed ad hoc models that manage to reproduce the general properties of stellar convective envelopes, including their thermal structure. Among them, we can cite: mixing length theory (MLT;

[★] Full Table A.1 is available at the CDS via anonymous ftp to cdsarc.cds.unistra.fr (130.79.128.5) or via <https://cdsarc.cds.unistra.fr/viz-bin/cat/J/A+A/687/A146>

Böhm-Vitense 1958) or full spectrum of turbulence models (FST; Canuto & Mazzitelli 1991, 1992; Canuto et al. 1996). The MLT reduces the spectral energy cascade of turbulence to a Dirac distribution, namely: the whole convective flux is carried by eddies of a single size. A hot parcel of matter, before dissolving into the surrounding medium, rises a certain distance (called the mixing length) $\ell = \alpha_{\text{MLT}} H_p$, where H_p is the pressure scale height and α_{MLT} is a free parameter. In deeper regions of the star, convective energy transport is very close to adiabatic, which is isentropic, the thermal structure is then independent of the convection formulation, and (in particular) of α_{MLT} , which is used in the models (Gough & Weiss 1976). However, when the adiabaticity assumption breaks down, typically in low-density near-surface layers in convective stellar envelopes, the value of α_{MLT} determines the degree of non adiabaticity in the model, namely: the entropy difference between the minimum of entropy occurring at the top of the convective envelope and the fully adiabatic interior. The FST represents an attempt to work around this dependence. It assumes a more complex form for the cascade, for instance, a Kolmogorov spectrum. Original formulations of the FST models do not introduce any free parameters and define the mixing length simply as the distance, z , to the nearest boundary of the convective zone. For the purposes of fine-tuning this approach, Canuto et al. (1996) introduced a dependence on the pressure scale height through a free parameter α_{FST} , intended to be small and to experiencing only small changes from a model to another.

However Ludwig et al. (1999) found, using the thermal structures obtained from 2D radiation hydrodynamics (RHD) stellar atmosphere models that α_{FST} does not remain small and varies by a similar amount as α_{MLT} over the HR diagram. Over recent decades, there have been many works highlighting the successes and caveats of MLT and FST models (e.g., Arnett et al. 2015; Trampedach 2010; Kupka & Muthsam 2017) and trying to improve or to go beyond them (e.g., Gough 1977; Grigahcène et al. 2005; Li 2012; Pasetto et al. 2014; Gabriel & Belkacem 2018).

Nonetheless, MLT and FST models remain widely used in stellar evolution codes. The most practical issue these models raise is: how do we choose the value of the free parameter α ? Overall, there are two ways to choose the α value. Firstly, through an optimization process, α can be tuned so that stellar models match a set of observable constraints for one or several stars. The most straightforward case is the widely used solar calibration, in which α is obtained by adjusting its value (together with the chemical composition) of models to reproduce the present-day global solar properties (e.g., Christensen-Dalsgaard 1982; Morel et al. 2000; Serenelli 2016; Vinyoles et al. 2017). The same α value is then used to compute stellar models. This assumes that a unique value of α is appropriate for all stars at all evolutionary stages but there is no reason a priori why this should be the case (Ireland & Browning 2018) and RHD simulations of convective stellar atmospheres consistently point to a variable α (Trampedach et al. 2014; Sonoi et al. 2019). A possible extension would be to use stars, or sets of them, as calibrators, ideally establishing a semi-empirical relation between α values and stellar fundamental parameters. There is, however, a lack of reliable calibrators for which the stellar parameters are known with high enough level of accuracy.

Secondly, a different approach is to consider α as a completely free parameter that can be optimized on star-by-star basis

Table 1. Summary of the different notations used to denote the specific entropy.

s_{ad}	Specific entropy of the adiabat in 1D stellar evolution models.
s_{RHD}	Specific entropy of the adiabat in 3D stellar atmosphere models. Entropy assigned to the inflow at the bottom.
s_{presc}	Value of the specific entropy of the adiabat, obtained from a prescription.
s_{bot}	Average specific entropy at the bottom layer of 3D simulations of stellar atmospheres.

to find best-fitted stellar models to a set of observables. This can be done using either stellar models that are computed as part of the optimization process or pre-computed grids of models with different α -values together with an interpolation algorithm (e.g., SPInS; Lebreton & Reese 2020) to find the best-fit model. This is actually not a calibration, as the resulting α values are not used as the basis for modeling other stars, only to obtain a best-fit model. However, such an approach offers little hope of improving on our formulations and understanding of convection.

To improve the use of MLT or FST modeling of convection, it is necessary to find a relationship between α and global stellar parameters (effective temperature, surface gravity, and chemical composition). This is the aim of α calibrations. Ludwig et al. (1999) used standard 1D stellar atmosphere models with a tunable α_{MLT} parameter (similar to what we have in stellar evolution models) and a grid of more realistic two-dimensional (2D) radiative-hydrodynamics stellar atmosphere models, to find a relation between α_{MLT} , T_{eff} , and $\log g$. More precisely, they adjusted the value of α_{MLT} so that the specific entropy of the adiabat of their envelope models matches the one of their 2D models. This work provided an expression for α_{MLT} as a function of effective temperature, T_{eff} , and surface gravity, g , of the models. Aiming at improving this pioneering study, similar work was carried out by Trampedach et al. (2014) using solar-metallicity 3D atmosphere simulations computed with the Nordlund & Stein (1990) code and later by Magic et al. (2015) with 3D STAGGER models with a metallicity of $[\text{Fe}/\text{H}] \in [-4.0; +0.5]$. More recently, Sonoi et al. (2019) used a grid of 3D stellar atmosphere models computed with CO⁵BOLD (i.e., COnservative COde for the COmputation of COmpressible COnvection in a BOx of L Dimensions, with $L = 2, 3$; Freytag et al. 2002, 2012; Wedemeyer et al. 2004) to obtain a representation of α as a function of global stellar parameters for MLT and FST modeling, as well as for different temperatures, T , versus the optical depth τ relations ($T(\tau)$ relations) used to describe the atmosphere. All those prescriptions for α are very easy to implement in a stellar evolution code. However, they are highly dependent on the physical ingredients, as it is shown by the important change of calibrated α values introduced by a change of $T(\tau)$ relation.

At the same time, Tanner et al. (2016) suggested directly adjusting the α parameter in stellar evolution calculations along the evolution such that the specific entropy of the model adiabat, s_{ad} , matches (at all times) s_{RHD} , the value obtained in RHD simulations for the same T_{eff} , $\log g$, and $[\text{Fe}/\text{H}]$ values (notations for the various specific entropies are summarised in Table 1). This has the advantage that the specific entropy has a well defined physical meaning and is linked to the surface properties of the star by means of the realistic RHD models. In this

¹ In the following, we use α equivalently for α_{MLT} or α_{FST} , except when explicitly mentioned otherwise.

way, α simply becomes a parameter used to ensure that stellar models reproduce the correct s_{RHD} , as inferred from RHD simulations, in a consistent way with the physical ingredients in the stellar evolution models; for instance, different $T(\tau)$ relations, use of MLT or FST, or any other prescription for convection. The idea of [Tanner et al. \(2016\)](#) was implemented in Yale Rotational stellar Evolution Code (YREC; [Demarque et al. 2008](#)) and tested on the Sun ([Spada et al. 2018](#)) and on the α Cen system ([Spada & Demarque 2019](#)). Finally, the method was extended to evolved stars ([Spada et al. 2021](#)).

The aim of the present work is to contribute to the general effort to develop more accurate stellar models. Modeling the atmosphere of solar-type stars is based on three main components: (1) the opacity profile in the atmosphere, (2) the $T(\tau)$ relation which defines the transition in the temperature gradient between optically thick and optically thin regions, and (3) the convection formalism. For a seismically reliable stellar model, these three components must be chosen with the utmost care. This work focuses on the last component by removing the main free parameters of the convection formalism. Specifically, in this work, we built on previous works to produce stellar models that no longer include the mixing length as a free parameter. To that purpose, we implemented the same method in the stellar evolution code Cesam2k20 (formerly known as CESTAM: Code d'Evolution Stellaire avec Transport, Adaptatif et Modulaire; [Morel 1997](#); [Morel & Lebreton 2008](#); [Marques et al. 2013](#)), as in the YREC code ([Demarque et al. 2008](#)). The previous works made it clear that a lot of care must be taken in the use of entropy prescription. In Sect. 2, we review the available prescriptions and in Sect. 3, we describe the implementation of entropy calibration in Cesam2k20. In Sect. 4 we describe the necessary improvements that should be added to the prescriptions. A discussion is presented in Sect. 5, followed by our conclusions in Sect. 6.

2. Available prescriptions for entropy at the bottom of the convective envelope

2.1. Stellar atmosphere codes

Before describing the entropy prescriptions, it is worth considering a few details on the 3D radiation-hydrodynamics model atmosphere codes that were used to compute the simulations they are based on. The atmosphere codes can solve hydrodynamic and radiative transfer equations. All simulations used hereafter do not include the magnetic field and therefore are referred to as RHD simulations. The codes rely on the “box-in-a-star” setup, where only the flow included in a box is simulated. This box covers an optical depth range of $-6 \leq \log_{10} \tau_{\text{Ross}} \leq 5$, which encompasses the superadiabatic region. This domain is small in comparison to stellar dimensions but large enough to resolve properties of the convective flow. Periodic boundary conditions are assumed at the vertical faces of the box and, at horizontal faces, the flow is free to enter and leave the box. The upper boundary condition can vary from one code to the other. At the bottom, the specific entropy of the inflowing material is fixed, which allows for the effective temperature of the resulting model to be constrained. The specific entropy of the inflow can be seen as the entropy of the adiabat. Because at the lower boundary, inflow and outflow do not have the same entropy, the horizontally average at this depth is not equal to the entropy of the adiabat.

In the following, we use s_{presc} to denote the specific entropy obtained from a prescription, s_{RHD} the specific entropy of the adiabat of an atmosphere model, and the word “entropy”, as

applied as a synonym for “specific entropy” (entropy per unit of mass). Up to now, three different mathematical formulations have been proposed to relate s_{RHD} to T_{eff} , $\log g$ and, optionally, $[\text{Fe}/\text{H}]$. While their general mathematical forms remain similar, they have been calibrated on different sets of models.

2.2. The functional form of [Ludwig et al. \(1999\)](#)

The original idea that models of surface convection could be used to derive constraints on the value of the entropy of the adiabat was first suggested by [Steffen \(1993\)](#). [Ludwig et al. \(1999\)](#) (hereafter [L99](#)) used a grid of 58 two-dimensional RHD atmosphere models with $4300 \text{ K} \leq T_{\text{eff}} \leq 7100 \text{ K}$, $2.54 \leq \log g \leq 4.74$ and a chemical composition close to that of the proto-Sun composition proposed by [Grevesse & Sauval \(1998\)](#) (hereafter [GS98](#)), with $Y = 0.28$ and $Z = 0.016$. The entropy of the adiabat is expressed as follows:

$$s_{\text{presc}}^{\text{L99}}(\tilde{T}, \tilde{g}) = a_0 + a_1 \tilde{T} + a_2 \tilde{g} + a_3 \exp(a_4 \tilde{T} + a_5 \tilde{g}), \quad (1)$$

where $\{a_i\}_{i=0}^5$ are constants determined using the grid of models, $\tilde{T} = (T_{\text{eff}} - 5770)/1000$ and $\tilde{g} = \log(g/27500)$.

2.3. The functional form of [Magic et al. \(2013\)](#)

[Magic et al. \(2013\)](#) (hereafter [M13](#)) used a grid (the STAGGER-grid) of 217 3D RHD atmosphere models computed with the STAGGER code ([Nordlund & Galsgaard 1995](#)), with $4000 \text{ K} \leq T_{\text{eff}} \leq 7000 \text{ K}$ in steps of 500 K, $1.5 \leq \log g \leq 5.0$ in steps of 0.5, and seven values of metallicity $[\text{Fe}/\text{H}] = +0.5, 0.0, -0.5, -1.0, -2.0, -3.0, -4.0$. The solar chemical composition follows the determination of [Asplund et al. \(2009\)](#) (hereafter [AGSS09](#)), for the present Sun. [M13](#) suggested that the entropy of the adiabat ought to be represented by the following function:

$$s_{\text{presc}}^{\text{M13}}(\tilde{T}, \tilde{g}, \tilde{z}) = P_2(a_i, \tilde{z}) + P_2(b_i, \tilde{z})\tilde{T} + P_2(c_i, \tilde{z})\tilde{g} + P_2(d_i, \tilde{z}) \exp(P_2(e_i, \tilde{z})\tilde{T} + P_2(f_i, \tilde{z})\tilde{g}), \quad (2)$$

and

$$P_2(a_i, \tilde{z}) = \sum_{i=0}^2 a_i \tilde{z}^i, \quad (3)$$

where the coefficients $\{a_i, b_i, c_i, d_i, e_i, f_i\}_{i=0}^2$ are constants, tuned to fit the entropy s_{RHD} of the STAGGER-grid models, $\tilde{T} = (T_{\text{eff}} - 5777)/1000$, $\tilde{g} = \log g - 4.44$, $\tilde{z} = [\text{Fe}/\text{H}]$. We stress that \tilde{z} must be computed with respect to an [AGSS09](#) chemical mixture. If one wants to use another chemical composition of reference, the coefficients in Eq. (2) should be recomputed. The expression proposed by [M13](#) is essentially the same as the one of [L99](#), except that the [M13](#) one takes into account the metallicity.

2.4. The functional form of [Tanner et al. \(2016\)](#)

Based on the same STAGGER-grid, [Tanner et al. \(2016\)](#) proposed a different form to represent s_{ad} as a function of T_{eff} , $\log g$ and $[\text{Fe}/\text{H}]$:

$$s_{\text{presc}}^{\text{T16}}(\tilde{T}, \tilde{g}; \tilde{z}) = s_0 + \beta \exp\left(\frac{A\tilde{T} + B\tilde{g} - x_0}{\tau_0}\right), \quad (4)$$

where $s_0, x_0, \beta, \tau_0, A$ and B are constants, calibrated for each subset of models with same metallicity; $\tilde{T} = \log T_{\text{eff}}$ and $\tilde{g} = \log g$.

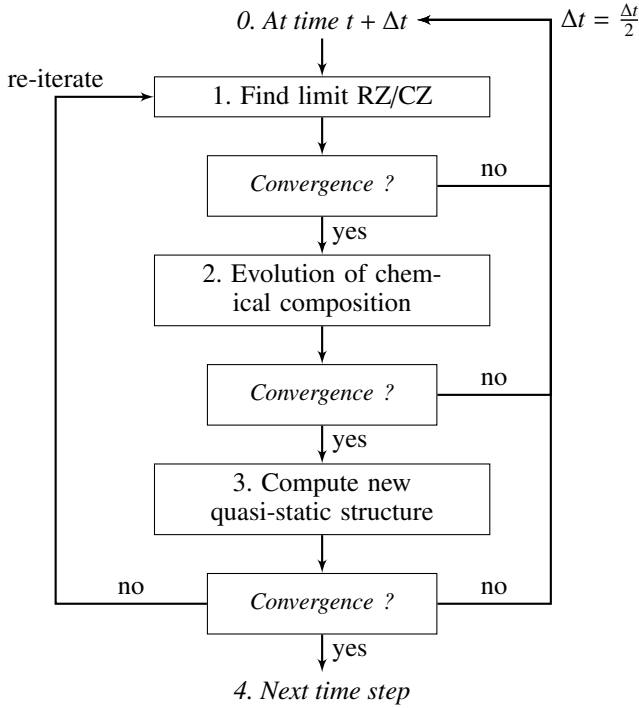


Fig. 1. Schematic representation of the step of a Newton-Raphson scheme followed by Cesam2k20 during the computation over a time step. The “problem” can be an unconverged solution, an interpolation outside EoS, or opacity table.

3. Implementation of entropy-calibration in Cesam2k20

3.1. The Cesam2k20 code

The Cesam2k20 code² implements the same numerical methods as its predecessors (for details, see Morel 1997 for CESAM, Morel & Lebreton 2008 for CESAM2k, and Marques et al. 2013 for CESTAM). The stellar structure equations are solved using a collocation method where solutions are represented as a linear combination of piecewise polynomials, projected over a B-Spline basis.

Cesam2k20 uses opacity tables either from the OPAL team (Rogers & Iglesias 1992; Iglesias & Rogers 1996) or OP tables from the Opacity Project (Seaton et al. 1994; Badnell et al. 2005; Seaton 2005, 2007). Several equations of state (EoS) are implemented and this work uses tables from the OPAL5 EoS (Rogers & Nayfonov 2002). Concerning the nuclear reaction rates, we used compilations from NACRE (Aikawa et al. 2006) (except LUNA – Broggini et al. 2018 – for the $^{14}\text{N}(p, \gamma)^{15}\text{O}$ reaction).

Regarding the mixing processes, rotation is neglected throughout this work. The effects of atomic diffusion are modeled with the formalism proposed by Michaud & Proffitt (1993). The convection is modeled with mixing-length theory formalism under the formulation of Henyey et al. (1965) taking into account the optical thickness of the convective bubble. We stress that the entropy-calibration method also works with more complex formalisms, such as the one developed by Canuto & Mazzitelli (1991) or by Canuto et al. (1996), assuming a Kolmogorov spectrum for the turbulent convective flow.

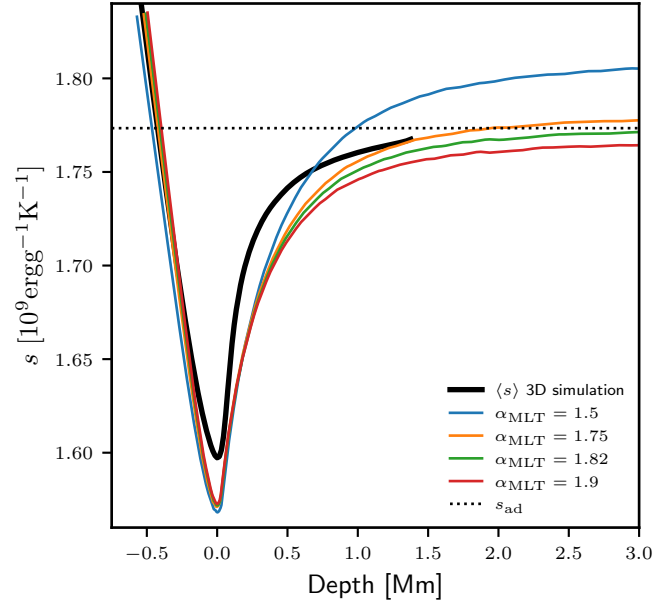


Fig. 2. Profile of specific entropy as a function of the depth in the model. The origin of the depth is taken as the location of then entropy minimum in the superadiabatic region. Average specific entropy profile of a 3D simulation is represented as a solid black line, and the entropy of the adiabat as a dash black line. The entropy of the adiabat is an input of the 3D model. Other curves are profiles of specific entropy obtained with 1D models computed with different values of $\alpha_{\text{MLT}} = 1.5$ (blue), 1.75 (orange), 1.82 (green), and 1.9 (red).

A schematic flowchart of steps followed by Cesam2k20 during the computation of a time step is represented in Fig. 1. At a given time step, Cesam2k20 iterates following a Newton-Raphson scheme to find the solution to the structure equations. More precisely, using the stellar structure determined at the previous time step, Cesam2k20 starts by finding the boundaries of the convective zones (CZ), then it evolves the chemical composition to the current time-step, then finally solving the stellar structure for that time. After each iteration, numerical convergence is assessed. If it is satisfied, the solution is accepted and new time step is computed. In general, this process is repeated several times to find an acceptable solution.

3.2. Scheme for entropy-calibration

As mentioned in the introduction, the value of α_{MLT} controls the adiabat of the convective envelope of the stellar model. This is easily seen in Fig. 2, where we present an average specific entropy profile extracted from a 3D atmosphere simulation and entropy profiles computed from 1D stellar structures. Despite the fact that 3D and 1D models give very different entropy profiles in the superadiabatic layers, they can converge in deeper layers to the same adiabat, providing we find the correct value of α_{MLT} .

The procedure of entropy calibrations requires an intermediate step between the computation of the new chemical composition and the resolution of the structure equations. This intermediate step is part of the Newton-Raphson iteration described in Eq. (1). In that step, we compute a preliminary version of the entropy $s_{\text{ad},1}$ and the mean molecular weight (the reasoning is discussed in Sect. 4.2) at the bottom of the convective envelope. These quantities are computed in an approximate stellar structure obtained with a given $\alpha_1 \equiv \alpha$, either from a previous time step or from a previous structure iteration. With the values

² <https://www.ias.u-psud.fr/cesam2k20/>

Table 2. Parameters of the SSM model.

Observational constraints		
$L [L_{\odot}]$	$T_{\text{eff}} [\text{K}]$	$(Z/X)_s$
1	5772	0.0231
Adjustable parameters		
Y_0	$(Z/X)_0$	α_{MLT}
$0.27334 \pm 4 \times 10^{-5}$	0.02651 ± 10^{-5}	$1.828 \pm 3 \times 10^{-3}$

Notes. Optimal parameters are obtained for the standard solar model SSM calibrated with the OSM stellar optimization tool.

of T_{eff} , $\log g$ and $[\text{Fe}/\text{H}]$ for this intermediate model, we evaluate the chosen functional representation of s_{ad} that provides the entropy $s_{\text{presc},1}$ of the adiabat that this model ought to have.

Then, the value of α_1 is modified by a small amount: $\alpha_2 = \alpha_1 + \delta\alpha$ (currently, $\delta\alpha = 0.001\alpha_1$). With α_2 , we call again the routine that solves for the stellar structure. Again, we compute $s_{\text{ad},2}$ at the bottom of CZ, and the value $s_{\text{presc},2}$ that it must take according to the chosen functional. We want to find the correction $\Delta\alpha$ applied to α_1 , such that at next iteration we have:

$$\delta s_1 + \frac{\delta s_2 - \delta s_1}{\delta\alpha} \Delta\alpha = 0, \quad (5)$$

with $\delta s_i = (s_{\text{presc},i} - s_{\text{ad},i})$, $i = 1, 2$. Then, $\Delta\alpha$ is:

$$\Delta\alpha = -\frac{\delta s_1}{\delta s_2 - \delta s_1} \delta\alpha. \quad (6)$$

In the Cesam2k20 input file, the fitting function proposed respectively by Ludwig et al. (1999), Magic et al. (2013), and Tanner et al. (2016) can be set through the option `nom_alpha` respectively equal to ‘entropy_ludwig99’, ‘entropy_magic13’, and ‘entropy_tanner16’.

3.3. Control of precision and performance

The agreement between s_{ad} and s_{presc} is controlled before the end of a time step. If the absolute value of the difference between the two exceeds $10^5 \text{ erg g}^{-1} \text{ K}^{-1}$, the solution is rejected and the computation of the current evolution step is started over again with a time step divided by two. This threshold was found by trial and error, and can be adjusted by the user. To give a comparison, in the STAGGER grid, the total variation of the entropy of the adiabat spans $1.6903 \times 10^9 \text{ erg g}^{-1} \text{ K}^{-1}$. Therefore, the criterion is almost always met at the first try and the time step rarely needs to be decreased.

In terms of performance, the use of entropy calibration significantly increases the computation time (by a factor 2 to 3). This is due to the fact that in order to evaluate the effect of a change of α on s_{ad} , one has to recompute the entire stellar structure with $\alpha + \delta\alpha$. YREC proceeds differently: the envelope and the rest of the star can be computed separately and this can provide the new adiabat $s_{\text{ad},2}$ much faster. In Cesam2k20, the interior and the envelope are solved together and the calculation cannot be decoupled.

4. Amendments to the entropy prescription

4.1. A first naive implementation

Once the entropy calibration procedure has been implemented into Cesam2k20, the next question is how the method impacts a

standard³ evolutionary track. To do so, we first calibrated a standard solar model (SSM) with a chemical composition following the determination of GS98 which is close to the one used in the 2D atmosphere grid that the L99 function was fitted on. The opacity and EoS are interpolated from the OPAL5 tables, convection is treated in the MLT framework, gravitational settling is modeled using the Michaud & Proffitt (1993) approximation and no rotation is assumed. The atmosphere is reproduced using the reduced $T(\tau)$ relation, known as the Hopf function, $q(\tau)$:

$$T(\tau) = T_{\text{eff}} \left[\frac{3}{4} (q(\tau) + \tau) \right]^{1/4}. \quad (7)$$

The Eddington $T(\tau)$ turns into a constant $q(\tau) = 2/3$. Other choices are discussed in Sect. 5.1.3.

This SSM is calibrated by adjusting the initial helium mass fraction Y_0 , initial metal to hydrogen ratio $(Z/X)_0$ content, and the parameter α_{MLT} to match (at a solar age) the luminosity, $1 L_{\odot}$ ($L_{\odot} = 3.828 \times 10^{33} \text{ erg s}^{-1} \text{ cm}^{-2}$; Mamajek et al. 2015), the effective temperature 5772 K, and the current surface $(Z/X)_s$ of the Sun with a maximum error of respectively 10^{-5} , 1 K and 10^{-5} . These parameters are adjusted using a Levenberg-Marquardt algorithm implemented in the Optimal Stellar Model program (OSM⁴; for detailed applications, see Castro et al. 2021). This package is currently interfaced with the Cesam2k20 code and finds the model that best matches a set of observational constraints (in that sense, it can be considered an ‘‘optimal’’ model). Best values found for adjustable parameters of the model are reproduced in Table 2. Once this optimal model has been obtained, we recompute the evolution by keeping everything identical, except that this time, the value of α_{MLT} is adjusted at each time step during its evolution to match the entropy prescribed by either L99, M13 or T16. These models are called entropy-calibrated models (ECMs).

The resulting evolutionary tracks are represented in Fig. 3, left panel. As can be seen, the four models follow very different paths. During the PMS, the locations of the respective Hayashi tracks are very far apart, while all models follow quite a similar Henyey track (Henyey et al. 1955). At a solar age, the shift in effective temperature can reach $\approx 130 \text{ K}$, with respect to the SSM. Models computed with M13 and T16 lead to similar results for evolved phases after the terminal age main sequence (TAMS), but they significantly differ from the one obtained with L99.

We see that ECMs are different from SSM, but are they better? The discrepancies obtained at a solar age are problematic: the Sun must have a definite entropy for its adiabat, however, all prescriptions give very different values for the same set of input values. With this first implementation, it would seem that entropy-calibration is not very robust for stellar models. Moreover, the shift observed on the RGB between the track computed with L99 on the one hand, and the tracks computed with M13 and T16 (on the other hand) may reveal large differences between the grids on which these functions are adjusted. These issues stress the fact that the prescriptions should not be used at face value (as given in the original papers), but at least corrected.

4.2. On the proper use of entropy prescriptions

In their works, Spada et al. (2018) and Spada & Demarque (2019) used the entropy given by T16’s mathematical formu-

³ In the following, ‘‘standard’’ means that α_{MLT} stays fixed along evolution.

⁴ <https://pypi.org/project/osm/>, developed in Python 2.7 by R. Samadi, adapted to Python 3 for the present work.

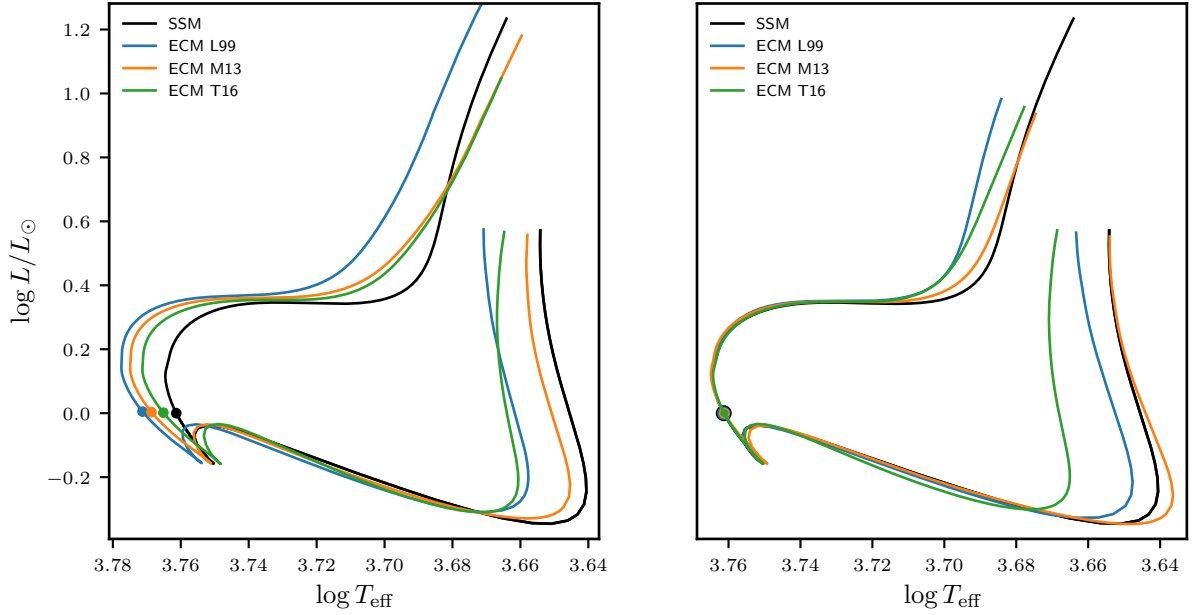


Fig. 3. HR diagram of a standard solar model (black line) and associated entropy-calibrated models shown in the left panel, with the uncorrected prescriptions: **L99** (blue), **M13** (orange), and **T16** (green) taken “as is”. Dots represent the location of each model at solar age. Right panel displays the HR diagram of a standard solar model (black line) and associated entropy-calibrated models with the prescriptions **L99** (blue), **M13** (orange), and **T16** (green) corrected using Eq. (11). Dots represent the location of each model at solar age.

lation corrected by an offset, ds . Indeed, the entropy is formally defined up to a constant. This constant can differ from one code to another because different EoS tables were used to generate either the atmosphere or the evolution models. Actually, entropies used in the STAGGER-grid to calibrate the coefficients in **M13** or **T16**’s functions were already defined using a constant offset to make possible their comparison to those in the **L99** models (see [Magic et al. 2013](#), p. 9).

To compute the offset ds that should be added to s_{presc} , we first need to compute a model with a standard (i.e., a fixed α) and then compare its adiabat to the one predicted by the chosen prescription. Of course, the standard model should use the same physics, and especially the same EoS, as the entropy calibrated model that will be computed.

The second effect that needs to be taken into account is the entropy variation due to a change of chemical composition. Strictly speaking, to conduct entropy calibration, we should always consider models with a chemical composition identical to the one used to fit the coefficients of the entropy representation. However, the chemical composition changes through the star’s life and this should be accounted for. In an ideal gas, neglecting the radiation pressure and assuming a polytropic EoS, one can write the entropy as (e.g., [Ireland & Browning 2018](#)):

$$s \simeq s_0 + \frac{N_A k_B}{\mu} \ln \left(\frac{T^{1/(\gamma-1)}}{\rho} \right), \quad (8)$$

where s_0 is a constant, μ is the mean molecular weight, T the temperature, ρ the density, and γ the adiabatic exponent. A change of chemical composition has its main impact through the mean molecular weight. Because the term s_0 in Eq. (8) is small compared to the second term⁵, the entropy of a given atmosphere model with a mean molecular weight changing from μ_{RHD} to μ

can be approximated as

$$s \simeq s_{\text{presc}} f_{\mu}, \quad \text{with } f_{\mu} = \frac{\mu_{\text{RHD}}}{\mu}. \quad (9)$$

This relation is valid only for a small change of μ .

This correcting factor was used in [Spada et al. \(2021\)](#) to account for significant chemical changes in evolved stellar phases but the offset correction, ds , was not used. We recommend, however, using both corrections. So, the next issue is to find out what form the correction should take. Coming back to the computation of the offset, ds , if the standard 1D model and the atmosphere model grid do not share the same chemical composition, then the offset simply computed as $ds = s_{\text{ad}} - s_{\text{presc}}$ does not only depend on a different choice of constant, but also on the different chemical mixtures. Instead, the offset should be adjusted as follows:

$$ds^{\text{presc}} = s_{\text{ad}}^{\text{std}} - s_{\text{presc}}(T_{\text{eff}}^{\text{std}}, \log g^{\text{std}}, [\text{Fe}/\text{H}]^{\text{std}}) f_{\mu}^{\text{std}}, \quad (10)$$

where “std” stands for quantities of the standard model, f_{μ}^{std} from Eq. (9) accounts for the change in mean molecular weight, from the entropy prescription (based on 3D RHD simulations) to the 1D standard model we want to apply it to. Then, the corrected entropy takes the form:

$$s_{\text{presc}}^{\text{corrected}} = s_{\text{presc}} f_{\mu} - ds^{\text{presc}}. \quad (11)$$

This correction is applied every time a prescription is evaluated because f_{μ} changes over the course of the evolution, for instance, as a result of gravitational settling or dredge up.

In order to see the impact of this correction on the evolutionary tracks, we repeated the experiment of Sect. 4.1 but this time the ECMs follow the corrected prescriptions. To compute the entropy offset of each entropy representation, we used SSM as our standard model. The values are given in the first line of Table 3. The evolutionary tracks of SSM and the new ECMs are presented in Fig. 3. Now, we find that all four tracks intersect at

⁵ This has been verified using data of the 3D atmosphere models grid CIFIST reproduced in Table A.1.

Table 3. Entropy offset Δs [$10^9 \text{ erg g}^{-1} \text{ K}^{-1}$] computed with the SSM model and for various prescriptions.

	Prescription		
	L99	M13	T16
With original coefficients:	0.011	0.042	0.021
With coefficients recalibrated on the CIFIST grid:	0.008	0.010	-0.001

Notes. The coefficients used in the prescriptions can either be taken from the original papers or re-calibrated on the CIFIST grid (see Appendix B; Ludwig et al. 2009).

solar age. However differences remain during the Hayashi tracks and on the RGB. The Hayashi track is a very short period of time (≈ 6 Myrs for the Sun), and its modeling is subject to a lot of uncertainties (e.g., Zwintz & Steindl 2022). The differences seen in this phase do not propagate to later evolutionary phases and do not affect the validity of the method. The differences for PMS and RGB models have two sources. First of all, these two regions correspond to locations where the three entropy functional representations have the higher inaccuracies, i.e. largest errors between fitting function and entropy of the adiabat coming from 3D models. Moreover, the entropy of the adiabat does not vary linearly with T_{eff} and $\log g$. The entropy fitting functions of Eqs. (1), (2), or (4), together with coefficients given in Appendix B, show that the entropy of the adiabat increases towards higher T_{eff} and lower $\log g$, but with decreasing slopes in the same direction. Therefore, a given uncertainty on the entropy coming from the fitting procedure, will imply a smaller change of location when the model is on the RGB region than when it is in the PMS. This explains why the discrepancies in the RGB region are smaller than around the Hayashi track. Still, some work remains to be done to improve the method in these evolutionary phases, for instance, by improving the analytic prescriptions in these regions of the Kiel diagram.

4.3. Recommendations on the use of prescriptions

In addition to the a posteriori correction of the prescribed entropy values, the 3D RHD simulation basis for the three formulations (L99, M13, and T16, i.e., a re-fit of M13) are not equivalent. The models used to calibrate the L99 law have been computed in 2D; therefore they may not be reliable enough to reproduce the entropy stratification of 3D models. Furthermore, this relation was obtained only for one (solar) metallicity. Therefore, the correction f_{μ} may not be valid because the difference between μ_{RHD} and μ from the entropy calibrated model may not be close enough to zero. For these reasons, we recommend not to use the L99 prescription (with any set of coefficients), except when computing non diffusive models, with a chemical composition similar to GS98. L99 is therefore not used in the remainder of this work.

Regarding the functional forms suggested by M13 and T16, it appears that they have not been originally calibrated against the “true” entropies of the adiabat of the STAGGER grid (i.e., the inputs of the models), but against s_{bot} the average entropy at the bottom of the simulation domain (s_{RHD} values are not publicly available). These simulations assume that at the bottom of the boxes, the adiabat is reached and $s_{\text{bot}} = s_{\text{RHD}}$. However, this

is only true for some models (see Table A.1) and it introduces significant disagreements between s_{RHD} and the value given by the prescription. On the contrary, in CO⁵BOLD models of the CIFIST grid, the true entropy of the adiabat is known exactly because it is an input of the simulation and corresponds to the entropy of the inflow at the bottom boundary of the simulation domain. This also mean that the value of s_{RHD} is not flawed by any boundary effect, which is not the case for the value of s_{bot} .

In order to perform a comparison of the various entropy representations (see Sect. 2) and to improve the quality of the entropy calibration, we re-calibrated the coefficients on a set of entropies of the adiabats obtained from 102 models of the CIFIST grid (Ludwig et al. 2009) computed with the CO⁵BOLD code (Freytag et al. 2002, 2012; Wedemeyer et al. 2004). Here, we use the true entropy of the adiabats, s_{RHD} (and not the one computed at the bottom of the simulated box). These data are available in Appendix A for models with solar metallicity and the coefficients of the functional forms are available in Appendix B. Figure 4 illustrates the impact of re-calibrating the M13 function, defined in Eq. (2). The left column shows $\Delta s^{\text{M13,o}} \equiv s_{\text{RHD}}^{\text{CIFIST}} - s_{\text{presc}}^{\text{M13,o}}$, namely, the entropy difference between the true entropy of the adiabat in the CIFIST grid and the original M13 formulation (“original” in the sense that the coefficients are taken from Magic et al. 2013). We note here that a constant has been added to the original M13 relation to compensate for the zero point offset between the entropy of the solar models in the CIFIST and STAGGER grids, as stated in Fig. 4 caption. The right column in Fig. 4 shows the differences $\Delta s^{\text{M13,r}} \equiv s_{\text{RHD}}^{\text{CIFIST}} - s_{\text{presc}}^{\text{M13,r}}$, in which $s_{\text{presc}}^{\text{M13,r}}$ is obtained after the M13 formulation has been re-calibrated in the CIFIST grid.

The distribution of Δs^{presc} for CIFIST models is also represented as a histogram in Fig. 5. We see that this new calibration reduces the range of entropy differences (from ≈ 0.05 to $\approx 0.03 \times 10^9 \text{ erg g}^{-1} \text{ K}^{-1}$) but the accuracy of M13 remains similar for the two grids. However, variations of Δs^{presc} across the Kiel diagram do not behave the same. In Fig. 4, left column, the entropy difference rises with decreasing T_{eff} , but in the right column, the larger differences are localized in the PMS region and around the RGB of stars more massive than the Sun. Those differences may arise from two effects: (1) the offset we added does not fully correct the discrepancies between the different EoS and reference solar composition (i.e., these differences do not induce a constant change across the Kiel diagram) and (2) in Fig. 4, left column, coefficients of Eq. (2) are calibrated on s_{bot} instead of s_{presc} .

The distributions of Δs^{presc} for the re-calibrated M13 and T16 formulations (Eq. (4)) are represented in Fig. 6. Both prescriptions hardly exceed an absolute error of $0.05 \times 10^9 \text{ erg g}^{-1} \text{ K}^{-1}$, which is less than 3% of the extremal variation of entropy across the CIFIST grid. This agreement is very satisfying. In addition, we note that the distribution of $\Delta s^{\text{M13,r}}$ is much more peaked around 0.0 than the one of $\Delta s^{\text{T16,r}}$. This comparison advocates for generalizing the use of the M13 prescription in entropy-calibrated modeling. From this point onwards, we only use the M13 prescription of adiabatic entropy.

5. Discussion

5.1. Possible sources of uncertainties

The accuracy of entropy calibration modeling can be affected by several factors. In the following section, we review them and estimate their impact on the determination of a star’s characteristics.

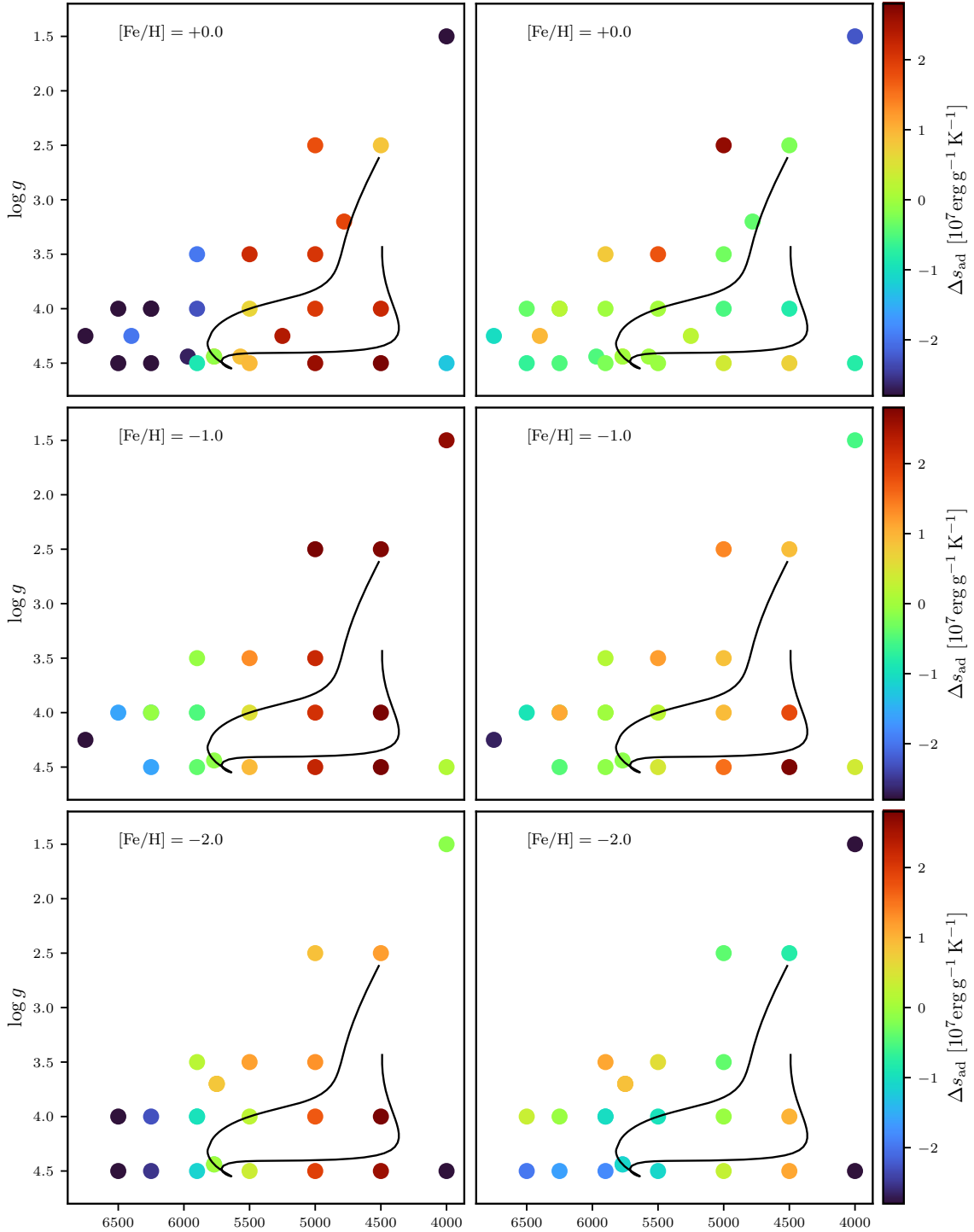


Fig. 4. Specific entropy differences in the adiabat $\Delta s^{\text{presc}} \equiv s_{\text{RHD}} - s_{\text{presc}}$ between values provided by the CIFIST grid (Ludwig et al. 2009), and values predicted by the M13 prescription, Eq. (2). Left column: original coefficients from Magic et al. (2013) are used. Here, entropy from the CIFIST grid has been corrected by an offset of $0.04794 \times 10^9 \text{ erg g}^{-1} \text{ K}^{-1}$ corresponding to the entropy difference between the solar atmosphere models in the CIFIST and STAGGER grids. Right column: same as left column but the coefficients involved in the M13 mathematical form have been re-calibrated with the CIFIST entropies. Rows: each row displays the models with a given metallicity: $[\text{Fe}/\text{H}] = +0.0$ (row 1), -1.0 (row 2) and -2.0 (row 3).

One source of uncertainty is the already mentioned gap between s_{RHD} and the corresponding value obtained from prescriptions. Accuracy can also be impacted by the methods used to determine the offset δs , along with the s_{ad} and μ values at the bottom of the convective envelope of the stellar evolution model.

Then, we should also examine what happens when a prescription is evaluated outside of its domain of definition. Finally, it could be expected that a different choice of $T(\tau)$ relation would change the location of an evolutionary track on the HRD or Kiel diagrams. These points are investigated in the following subsections.

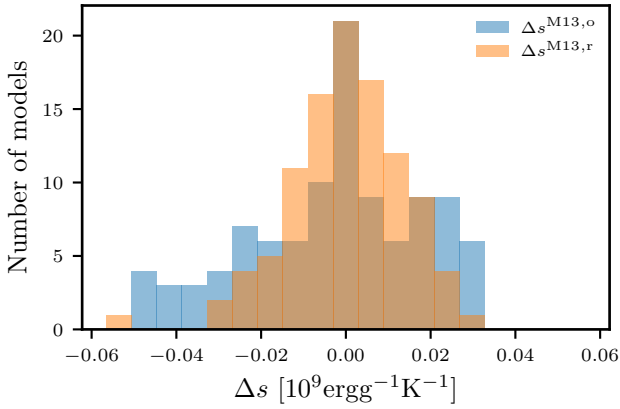


Fig. 5. Histogram showing the distribution of the entropy errors Δs^{presc} for the **M13** prescription. The data plotted are the same as shown in Fig. 4, for all metallicities. In blue is shown the $\Delta s^{\text{M13,o}}$ distribution corresponding to differences $s_{\text{RHD}}^{\text{CIFIST}} - s_{\text{presc}}^{\text{M13,o}}$, where $s_{\text{presc}}^{\text{M13,o}}$ represents the entropies predicted by the **M13** prescription with original coefficients (taken from [Magic et al. 2013](#)). The $\Delta s^{\text{M13,r}}$ distribution is shown in orange, corresponding to $s_{\text{RHD}}^{\text{CIFIST}} - s_{\text{presc}}^{\text{M13,r}}$, where this last term represents the entropies obtained from the **M13** prescription re-calibrated on the CIFIST grid (i.e., on the true entropies of the adiabats).

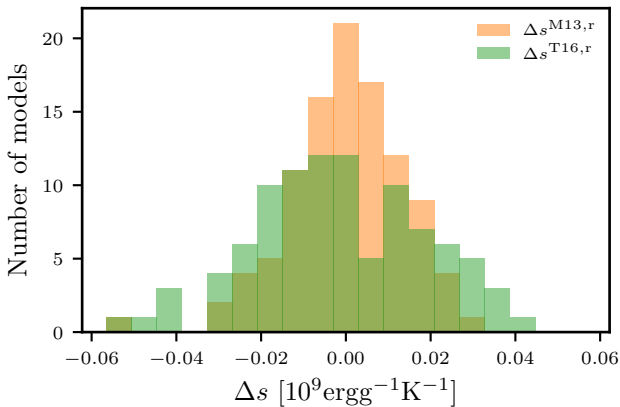


Fig. 6. Histogram showing the distribution of the entropy errors between s_{RHD} and s_{presc} for **M13** and **T16** entropy representations. Both functions have been re-calibrated on entropies from the CIFIST grid.

5.1.1. Determination of μ and s_{ad} in 1D models

Several factors can affect the accuracy of entropy-calibration modeling. The first one comes from the way we determine the mean molecular weight μ (in order to compute f_{μ}) and the entropy of the adiabat in the 1D model.

The mean molecular weight depends on the relative abundance and on the ionization rates of each element. Strictly speaking, the mean molecular weights in the 3D simulations and in the 1D models should be compared assuming well-defined ionization states between the two chemical compositions. The proper way to compute $\mu_{1/3D}$ would be to use the same EoS in the 1D and 3D code, and evaluate them for a given state (given pressure, temperature and chemical composition). However, we want to be able to use any EoS in the 1D code, possibly different from the one used in the 3D simulations. Moreover, not only is the chemical composition different in the 1D and 3D simulations, but also the elements considered are not the same. To get around these issues, we can assume that at the bottom of the convective zone, the material is fully ionized. This is not exactly

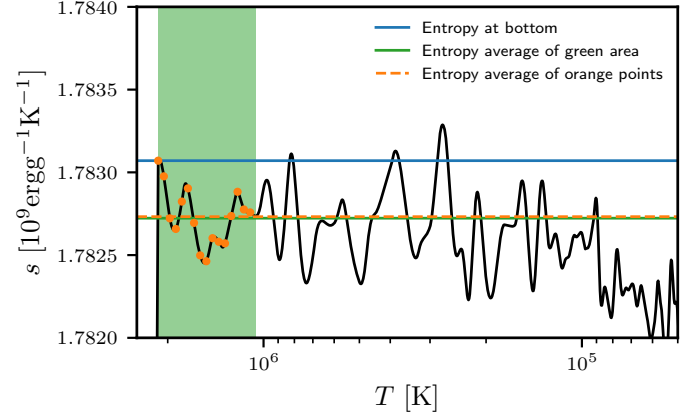


Fig. 7. Entropy as a function of temperature (black line) in a sub-region of the convective envelope of SSM. The value of entropy at the limit between radiative and convective zone is represented as a blue line. The average of entropy in the green shaded area is represented as a green line and the average of entropy at orange points is represented as a dashed orange line. The green area has a width of 10% of the total width of the CZ (roughly $2/3$ of the mass of the CZ) and the orange points are located only every ten layers in the green area.

true, but it has the merit of considering the same ionization state for the two mixtures. In this case, μ can be approximated as $\mu^{-1} = 2X + 3Y/4 + Z/2$ where X , Y and Z are respectively the hydrogen, helium and metal abundance. The error on f_{μ} induced by this approximation is below 1% for stars with $M < 1.5M_{\odot}$. We could relax this approximation by assuming that the ionization rates are the same in the evolutionary and in the atmosphere models. This would allow us to consider each element individually, instead of a dependence in X , Y , and Z . However, for this to work, it would require us to consider exactly the same chemical elements in both simulations, which would involve too much constraint from a modeling point of view. In any case, this would modify the value of μ by an almost constant factor, which would be taken into account in f_{μ} , as defined in Eq. (9).

With respect to the determination of the entropy of the adiabat in the 1D model, we compared the three available methods. In the first one, we defined $s_{\text{ad}} \equiv s_{\text{ad}}^{\text{bot}}$ as the value of the entropy exactly at the transition between the radiative zone and the convective envelope. Another method would be to define $s_{\text{ad}} \equiv s_{\text{ad}}^{\text{av}}$ as the average of the entropy in a sub-region of the adiabatic region. This second method would require us to interpolate in the EoS table for each layer in the sub-region, which is time-consuming. To improve its efficiency, we introduce instead a third definition, whereby $s_{\text{ad}} \equiv s_{\text{ad}}^{\text{av}10}$ is the average of the entropy computed for every ten layers in the sub-region only.

Figure 7 represents the entropy profile in the adiabatic region of SSM. We stress that the figure is extremely zoomed-in and what seems to be large fluctuations is actually nearly constant. Starting from $T \simeq 10^5$ [K] to the right, we see a decrease of entropy, corresponding to the beginning of the superadiabatic region. The value of $s_{\text{ad}}^{\text{bot}}$ is represented as the blue line. The green shaded area corresponds to the sub-regions where $s_{\text{ad}}^{\text{av}}$ and $s_{\text{ad}}^{\text{av}10}$ have been defined. The green region represents 10% of the radial extension of the convective envelope and $\simeq 66\%$ of its total mass. This extent is chosen arbitrarily but ensures that, for every model in the domain of definition of the CIFIST grid, we only enclose a sub-region of the adiabatic region. Clearly, $s_{\text{ad}}^{\text{av}}$ and $s_{\text{ad}}^{\text{av}10}$ are nearly equal. So, we consider why the use of

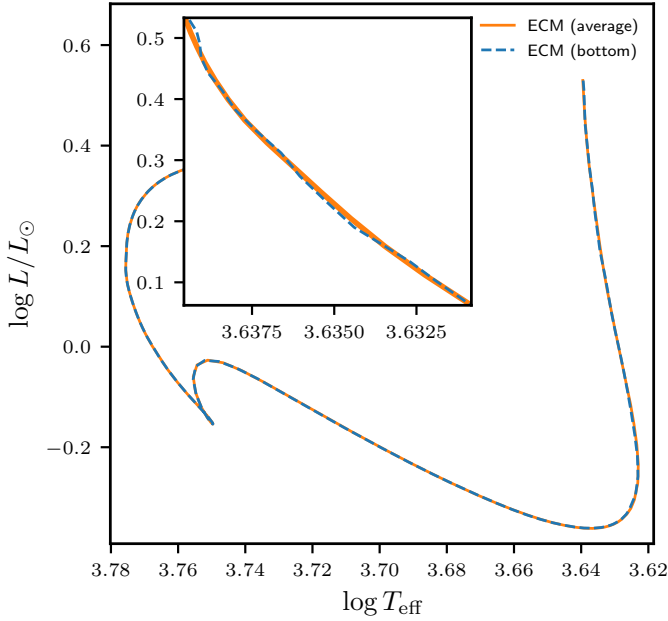


Fig. 8. HR diagram of the same entropy-calibrated model, using the M13 prescription, but different methods for evaluating the entropy of the adiabat in the 1D structure model: $s_{\text{ad}}^{\text{av}10}$ in orange and $s_{\text{ad}}^{\text{bot}}$ in blue. The subplot shows a zoom of the Hayashi track.

$s_{\text{ad}}^{\text{av}10}$ instead of $s_{\text{ad}}^{\text{bot}}$ would give significantly different results and discuss these possibilities below.

We implemented both definitions in Cesam2k20 and computed two evolutionary tracks using the M13 prescription. Tracks are shown in a HR diagram in Fig. 8. In terms of global quantities, the two definitions give similar results. However, as stressed by the subplot showing a zoom of the Hayashi track, using $s_{\text{ad}}^{\text{bot}}$ gives a more irregular path than $s_{\text{ad}}^{\text{av}10}$, whereas using $s_{\text{ad}}^{\text{av}10}$ actually significantly increases the numerical stability of the method.

5.1.2. Constancy of the entropy offset, d_s

The way we computed the offset, d_s , immediately brings on the concern that it may vary across HR diagram. As said before, the offset d_s corrects for a systematic difference in the entropy of the 3D atmosphere grid and the one computed in the stellar evolution model. This offset is the sum of the entropy zero-point for that EoS, and an ad-hoc constant for it to match the L99 entropy. The offset should be estimated with a standard solar model (standard in the sense that α_{MLT} stays fixed along evolution), which provides us with the most accurate constraint one can obtain on the entropy of the adiabat.

In order to verify that d_s is almost constant over the HR diagram, we computed entropy for three different EoS (numerical or analytical), for a given chemical composition, and conditions of pressure and temperature that cover the ones found at the bottom of convective envelope of stellar models with a mass between 0.8 to 1.3 M_{\odot} . The three EoS are the OPAL5 equation of state, the simple EoS proposed by Wolf (1983, hereafter noted as W83⁶) and FreeEOS (Irwin 2012). OPAL5 is the EoS table used to compute Cesam2k20 models. The simple ana-

lytic EoS proposed by W83 is the one used by CO⁵BOLD. Here, we implemented a simplified version which accounts only for H^+ , He^+ , He^{++} , e^- and a representative metal (case f. in W83). FreeEOS is also an analytic EoS but much more detailed. We followed FreeEOS’s recommendation and adopted the version suitable for stellar interiors that also reproduces the MHD equation of state (Mihalas et al. 1988) used by the STAGGER code. This version includes 20 elements with 295 ionization states. The entropy differences between FreeEOS (resp. W83) and OPAL5 are presented in a $\log_{10} p - \log_{10} T$ plane in left (resp. right) panel of Fig. 9. Extremal variation in this interval is on the order of $10^7 \text{ erg g}^{-1} \text{ K}^{-1}$. This is negligible compared to, for instance, the extremal variation of entropy across the CIFIST grid ($1.69 \times 10^9 \text{ erg g}^{-1} \text{ K}^{-1}$). In addition, these extremal variations are reduced if we exclude top left (resp. bottom-right) corner that corresponds to regions of low p – high T (resp. high p – low T), which are never found at the bottom of the convective envelope of our models. Therefore, the assumption that offset, d_s , is constant over the HR diagram holds to good accuracy.

5.1.3. Impact of the $T(\tau)$ relation

The previous attempts to prescribe a value for α_{MLT} have in common the weakness that the α prescription depends on the kind of $T(\tau)$ relation chosen to reproduce the atmosphere. This issue disappears with entropy calibration modeling, as shown in Spada et al. (2018) for the Sun and in Spada et al. (2021) for MS and RGB stars. The reason is that the entropy of the adiabat depends on the characteristics of the adiabatic layers, not on how the atmosphere is computed. The adiabat of the convective envelope of a star depends solely on the star’s location in the Kiel diagram and its metallicity. If the $T(\tau)$ relation changes, the value of α_{MLT} will adapt, but the location on the Kiel diagram is the same. To verify that this property remains true with our procedure, we computed four identical solar entropy calibrated models except for the $T(\tau)$ relation that follows either an Eddington law, a fit of Vernazza et al. (1981)’s Model C of the quiet Sun, the Krishna Swamy (1966) relation, or a $T(\tau)$ relation recently proposed by Ball (2021).

A change in the $T(\tau)$ relation, at a fixed α_{MLT} , changes global quantities such as T_{eff} . However, here, we are adjusting α_{MLT} to match a given entropy of the adiabat, therefore changing the $T(\tau)$ relation does not imply a change of location in the HR diagram. In the top panels of Fig. 10, we represent the evolution with age, starting from PMS, of the relative L , R , T_{eff} and the absolute s_{ad} differences for models computed with various $T(\tau)$ relations, with respect to a standard model computed with the Eddington $T(\tau)$ relation. It clearly shows the invariance of such global quantities when a different $T(\tau)$ law is chosen. On the bottom left panel, we show α_{MLT} along evolution for the same models. Since, T_{eff} , $\log g$ and the metallicity are the same from one model to the other, s_{ad} also remains constant. The structure of superadiabatic layers at final age (4570 Myrs) is shown in the bottom panel of Fig. 10. It stresses that $T(\tau)$ relations have an important impact on the modeling of the shallowest layers but the structures are all identical when deep enough in the star. The deepest layer shown in the figure has a pressure of 10^7 dyn cm^{-2} , which is well above the bottom of the convective envelope, located around a pressure of $10^{13} \text{ dyn cm}^{-2}$. It highlights the fact that although the choice of $T(\tau)$ relation does not impact global properties of the model, this choice is still crucial for obtaining realistic models for asteroseismology.

⁶ Notice the typo in the summation part of their Eq. (16): one should not read m_e , the mass of an electron, but m_i the mass of the considered element.

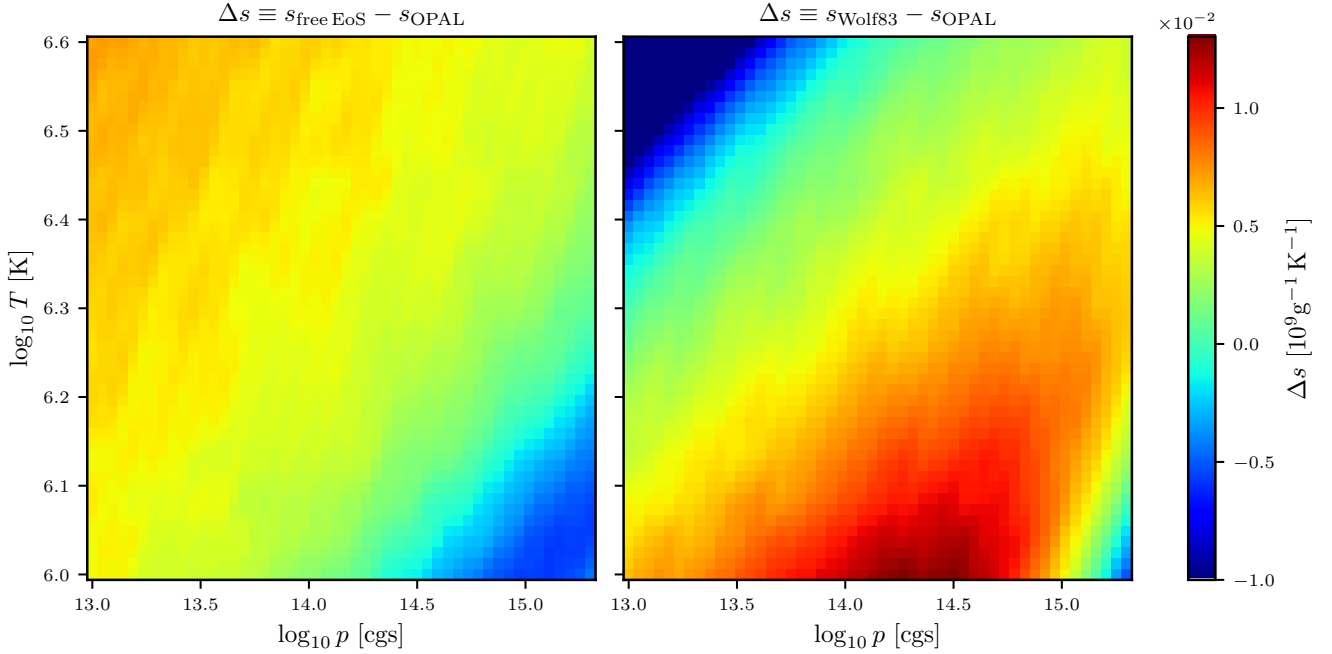


Fig. 9. Entropy differences between two EoS, in the $\log_{10} p - \log_{10} T$ plane. The colour codes for the entropy difference. Left panel: entropy differences between free EoS and OPAL5. Right panel: entropy differences between EoS of [Wolf \(1983\)](#) and OPAL5.

The examination of the bottom left panel of Fig. 10, reveals a minimum in the α_{MLT} variation with age between 2 and 6 Myrs, when models are on their Hayashi track. It is not possible to relate this behavior to previous calibrations of α_{MLT} of [Ludwig et al. \(1999\)](#) or [Magic et al. \(2015\)](#), but only with the one of [Trampedach et al. \(2014\)](#) or [Sonoii et al. \(2019\)](#) for what they called “Eddington + MLT(BV)”. This last work calibrated α_{MLT} by matching 1D CESTAM structure models to 3D CIFIST atmosphere models. CESTAM (and also its successor Cesam2k20) implements two versions of the MLT that assume either a linear temperature distribution in the convective bubble (that [Sonoii et al. 2019](#) called “BV”) or a distribution that obeys a diffusion equation, as suggested by [Henyey et al. \(1965\)](#). The BV version of MLT is also the version used in the work of [Trampedach et al. \(2014\)](#) and in this work, which explains the agreement.

5.1.4. Interpolation outside the prescription’s domain of definition

It may happen that at some stage of the evolution, a model falls outside the domain of definition of the prescription, which is the domain covered by the CIFIST grid in this work. For instance, when starting from the PMS, a solar model spends its first 1 Myrs slightly outside this domain.

Every time a given prescription is evaluated, Cesam2k20 checks if the model (its $(T_{\text{eff}}, \log g, [\text{Fe}/\text{H}])$) is inside the convex hull of the set of 3D models. By default, Cesam2k20 only sends a warning (this behavior can be changed by the user), and proceeds to evaluate the prescription. It must be noted that if prescriptions are evaluated outside of the domain, on the hot side, the resulting entropy can diverge because of the positive term in the exponential of Eqs. (1) and (2)). On the cool side, the exponential should remain small and the entropy does not diverge (which does not mean its value is reliable). The only solution to circumvent this issue is to extend the range of definition of the prescriptions. This work is currently in progress.

5.2. Comparison with YREC: The α Cen system

As a final test, we compared the results of modeling the α Cen system computed with standard or entropy-calibrated modeling, to the models presented in [Spada & Demarque \(2019, hereafter SD19\)](#). The microphysics of the models is kept as close as possible to the one used in YREC models: the MLT treatment of convection, Eddington atmosphere, and element abundance ratios following the [GS98](#) chemical composition and only gravitational settling is taken into account in the diffusive treatment. [SD19](#) used the YREC code to obtain best fitting models of α Cen A and B, using a Markov chain Monte Carlo (MCMC) method (more details given in [SD19](#)). The use of the MCMC method allowed the authors to provide reliable uncertainties to their best estimated parameters. To speed up the process, evolution models were started at the ZAMS.

In the present work, we found optimal Cesam2k20 models for the α Cen system using the OSM program. In our calibrations, we used the same observational constraints as the one used by [SD19](#), summarized in Table 4 and we adjusted the age of the model, its initial helium content Y_0 , its initial $(Z/X)_0$ ratio and (only for the standard models) the α_{MLT} , so that we match the observed radius, luminosity, and current $(Z/X)_s$. In ECMs, of course, α_{MLT} is varying along evolution, following the change of entropy. The mass is not a free parameter: we used the value reported in Table 4. In parallel, using the same physical ingredients, we calibrated a solar model to be the standard model for the computation of entropy offset d_s . Using the [M13](#) prescription yields $d_s = 0.0264 \times 10^9 \text{ erg g}^{-1} \text{ K}^{-1}$ and we used this value to compute entropy-calibrated models by tuning the same parameters as before (except for α_{MLT}). The Levenberg-Marquardt scheme implemented in OSM cannot yield reliable uncertainties and therefore, we give our best estimated parameters, without uncertainties. The squared-distance χ^2 between the best Cesam2k20 model and the observational constraints is computed with:

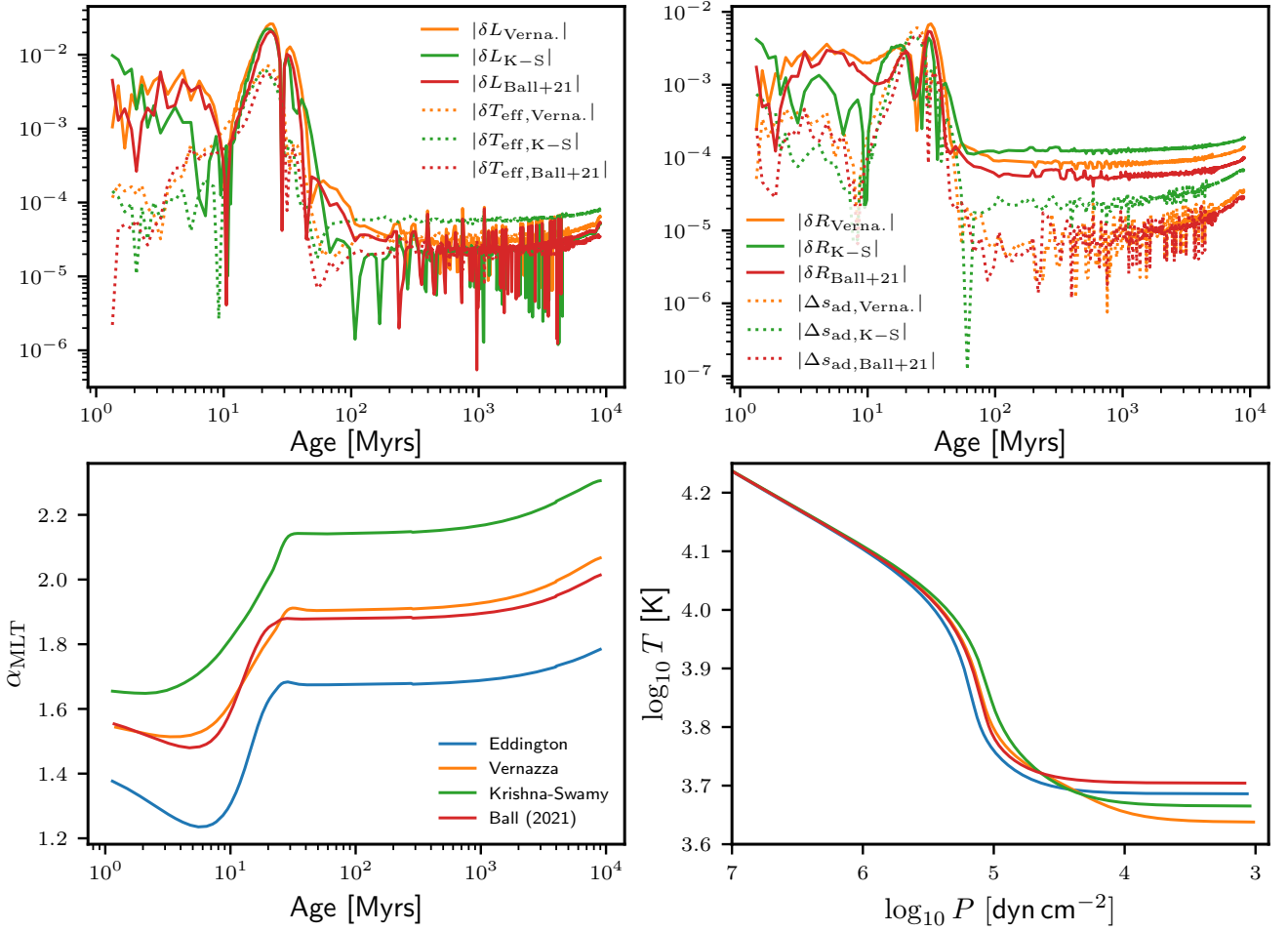


Fig. 10. Comparison of luminosity (solid lines) and effective temperature (dashed lines) for four models computed with different choices of $T(\tau)$ relation (top-left). The displayed quantity is the relative difference of L or T_{eff} , with respect to model computed with the Eddington $T(\tau)$ relation: $\delta X \equiv (X_{T(\tau)} - X_{\text{Edding.}})/X_{\text{Edding.}}$. Top right panel is the same as top left panel but, instead of $X \equiv s_{\text{ad}}$, we display the absolute difference. Indeed, with s_{ad} being defined up to a constant, the relative difference would have no meaning. Also, Δs^{presc} is in the units: $10^9 \text{ erg g}^{-1} \text{ K}^{-1}$. Bottom left panel: evolution of α_{MLT} for different $T(\tau)$ relations: Eddington (blue), Vernazza et al. (1981) Model-C (orange), Krishna Swamy (1966) (green), and Ball (2021) (red). Bottom right panel: temperature as a function of the pressure for the same four final P models at solar age.

$$\chi^2 = \sum_i \left(\frac{X_{\text{obs}}^i - X_{\text{model}}^i}{\sigma_i} \right)^2, \quad (12)$$

where X_{obs}^i are the observational constraints, σ_i their uncertainties, and X_{model}^i are the values of the same parameters in the best model. The MCMC method used by SD19 explores a wide parameter space, and then requires the computation of more models than OSM. Therefore, we could afford to start our Cesam2k20 models from the PMS.

The optimal parameters obtained with YREC and Cesam2k20 are presented in Table 4. Top panels of Fig. 11 show the tracks of all models in the radius – luminosity plane with the measured location of α Cen A and B. Models taken from SD19 show slightly worse agreement with observational constraints than models computed for the present work. This is only a consequence of the optimization scheme: since the MCMC method explores all the parameter space, a very large number of models need to be computed and obtaining a very precise solution is extremely time consuming. Alternatively, OSM requires to compute only a small number of models and the optimal solution is guided to a (possibly local) minimum.

The differences obtained between YREC and Cesam2k20 models can have a number of causes. First of all, the numerical schemes differ. The convective envelope is treated separately from the core in YREC, which allows it to tune α by recomputing only the envelope structure. This is not possible with Cesam2k20 and we have to recompute the complete structure every time α is changed. Also, in the case of ECM models, different entropy prescriptions are used and the way of correcting them differs. SD19 also made use of the T16 function with the coefficients taken from the original paper while we used the M13 function with coefficients re-calibrated on the CIFIST grid. In addition, it must be stressed that the model of α Cen B has its Hayashi track outside the domain of definition of M13, which was not the case for models in SD19, because their tracks started from the ZAMS. However, it represents a very short time of the life of α Cen B and should not have significantly impacted its evolution. Finally, SD19 corrected the prescribed entropy using only the entropy offset while Cesam2k20 also corrected for differences in the chemical composition. These discrepancies translate in quite different evolution of α_{MLT} between models with similar physics. However, the general trend remains similar: α_{MLT} increases when s_{ad} increases. The variations of entropy can easily be understood with the approximation that $s_{\text{ad}} \propto \ln T^{3/2}/\rho$.

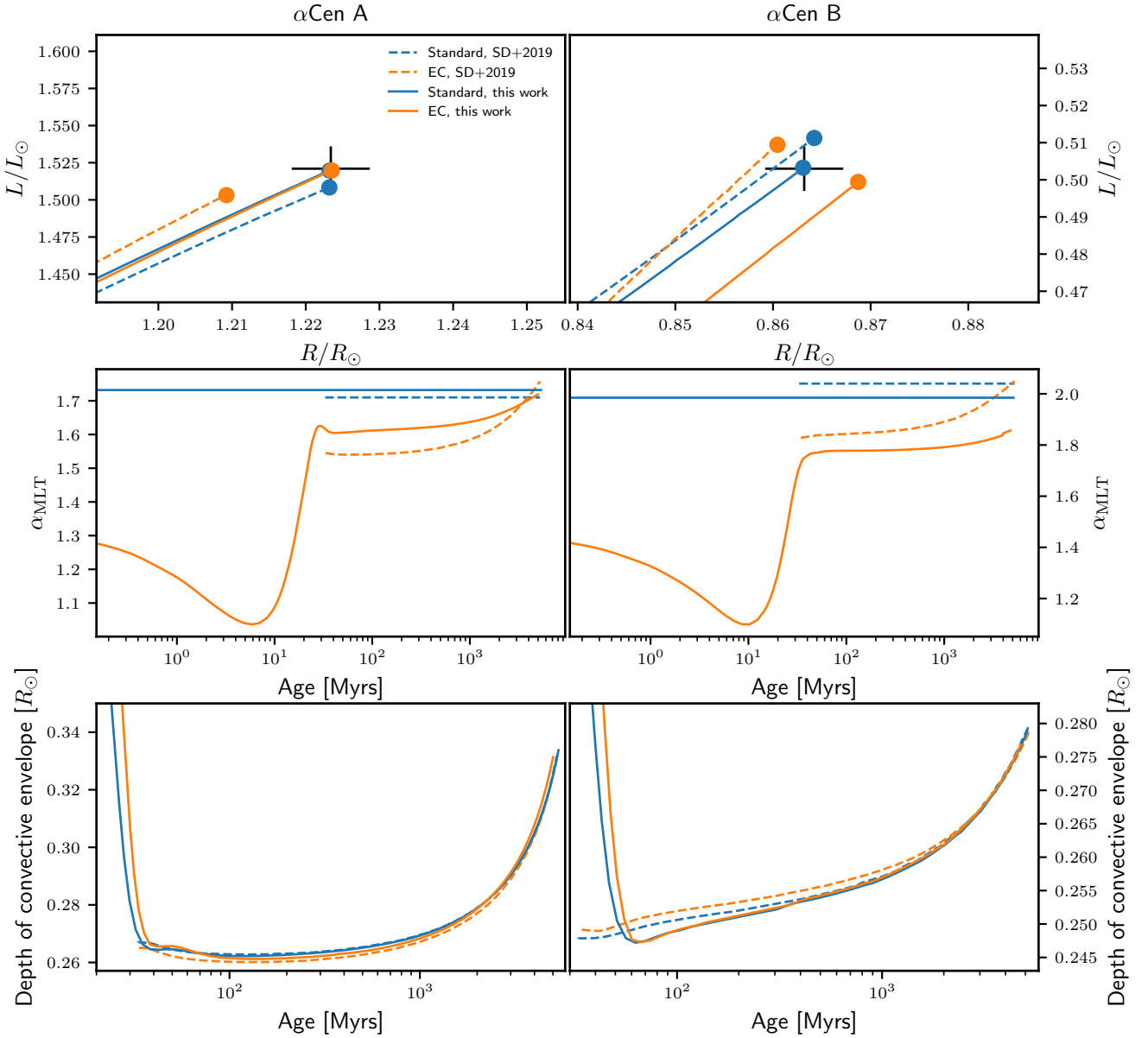


Fig. 11. Evolutionary tracks (top) for α Cen A (left panel) and α Cen B (right panel) on the $R - L$ plane. Tracks computed in the present work are represented as solid lines, while the one from Spada & Demarque (2019) are drawn as dashed lines. Standard models are in blue and EC models in orange. Evolution of α_{MLT} as a function of age (middle) for the same models. Evolution of the convective envelope depth (bottom) along the evolution.

With the virial theorem, it follows that $T \propto R^{-1}$ and $\rho \propto R^{-3}$, with R being the total radius of the star. Therefore, during expansion phases, when R increases, s_{ad} increases, and the inverse is true during the contraction phases.

The bottom panels of Fig. 11 display the variation of the depth of the convective envelope d_{cz} for all models of the α Cen system. Once again, it reveals a very good agreement between YREC and Cesam2k20 models. Whereas α_{MLT} depends a lot on the physical ingredients and numerical methods used in the code and does not have a precise physical meaning, d_{cz} is a measurable quantity. Having such an agreement for d_{cz} between the two codes advocates for the reliability of the method. Moreover, the values of d_{cz} predicted for a standard and EC models are close, but not identical. For α Cen B (the less massive), the two values are almost equal during the main sequence, while in the case of α Cen A (the more massive), the ECM predicts a shallower CZ

at the ZAMS and deeper at the present age of α Cen A. If confirmed, such pattern should have repercussions on the transport of chemicals and angular momentum.

6. Conclusion

In this work, we present a comprehensive description of the entropy-calibrated modeling, along with details on how to properly implement and apply it. The entropy-calibration consists in adjusting the value of the convection parameter, α , so that the entropy of the adiabat in a stellar structure model matches the one predicted by an entropy prescription.

We studied and compared the accuracy of the three available prescriptions for the entropy of the adiabat suggested in Ludwig et al. (1999), Magic et al. (2013) and Tanner et al. (2016). Those functions are adjusted to fit the entropy of the adi-

Table 4. Observable constraints and optimal parameters derived for the α Cen system.

Parameter	Obs. constraints	Standard		Entropy-calib.	
		YREC (SD19)	Cesam2k20	YREC (SD19)	Cesam2k20
α Cen A					
Mass (M_{\odot})	1.1055 ± 0.004 (1)	–	–	–	–
Radius (R_{\odot})	1.2234 ± 0.0053 (1)	1.2235	1.2233	1.2097	1.2235
Luminosity (L_{\odot})	1.521 ± 0.015 (1)	1.5091	1.5197	1.5045	1.5198
(Z/X) _s	0.039 ± 0.006 (2,3)	0.0315	0.0327	0.0347	0.0361
Age (Gyrs)	–	5.243	5.35	5.243	5.00
α_{MLT}	–	1.71	1.732	1.758 ^(*)	1.719 ^(*)
Y_0	–	0.2785	0.2781	0.2785	0.2898
Z_0	–	0.027	0.0275	0.027	0.0298
χ^2	–	2.21	1.12	8.38	0.24
α Cen B					
Mass (M_{\odot})	0.9373 ± 0.003 (1)	–	–	–	–
Radius (R_{\odot})	0.8632 ± 0.004 (1)	0.8646	0.8631	0.8608	0.8687
Luminosity (L_{\odot})	0.503 ± 0.006 (1)	0.5122	0.5032	0.5103	0.4994
(Z/X) _s	0.039 ± 0.006 (2,3)	0.0346	0.0357	0.0364	0.0339
Age (Myrs)	–	5.263	5.11	5.263	4.75
α_{MLT}	–	2.04	1.985	2.05 ^(*)	1.86 ^(*)
Y_0	–	0.2668	0.2666	0.2668	0.2640
Z_0	–	0.0273	0.0278	0.0273	0.0263
χ^2	–	3.00	0.30	2.035	2.99

Notes. ^(*)The values of α_{MLT} are the final values of the model, as for ECM, α_{MLT} is varying through evolution, as shown in Fig. 11, middle row.
References. (1): Kervella et al. (2017); (2): Thoul et al. (2003); (3): Porto de Mello et al. (2008).

abat in stellar atmosphere simulations, as a function of T_{eff} , $\log g$, and [Fe/H] in a grid of such simulations. Our study points out the fact that the M13 fit was based on the entropy, averaged over both the adiabatic up-flows and the cooled down-flows, at the bottom boundary of their RHD simulations – instead of on the entropy of the asymptotic adiabat. T16 is a refitting of those same, flawed entropies, introducing a bias on the prescribed entropy value. After the re-calibration of the entropy functional representations using the correct set of entropies of the adiabat obtained from the CIFIST grid (Ludwig et al. 2009), we show that the functional form that most accurately reproduces the RHD results and, therefore, the one we recommend is the formulation proposed by Magic et al. (2013) with coefficients given in Table B.2. We have made the quantities extracted from the CIFIST grid that have been used in this work publicly available. The coefficients of the fitting functions are also given in Appendix B, which extends the small list of entropy prescriptions available to the stellar modeling community.

We note that not only must the prescription be carefully chosen, but its use must be correctly set up. Spada et al. (2018) first proposed to correct the prescribed entropy with an offset that corresponds to a different choice of entropy integration constant between different EoS tables. In contrast, Spada et al. (2021) proposed to only correct the prescription for the influence of a different chemical composition between the one used in the evolution and in the atmosphere model. We demonstrate in this work that both corrections must be used. In order to compute entropy-calibrated models, the steps are as follows:

1. Compute a standard (i.e. not with entropy-calibration) model of reference. We advocate using a solar model as the reference because it is the star with the best observational constraints available. This standard model must implement exactly the same physical ingredients as the

future EC model (same convection model, same opacity table, etc.).

2. Compute the entropy offset between the standard model and the value it should have according to a given prescription. Of course this value changes for a different mathematical formulation and a different set of coefficients.
3. Knowing the offset value, the EC model can be computed. Each time the entropy prescription is evaluated, the resulting value must be corrected with the offset and for the different chemical composition.

This work also describes the traps that should be avoided when implementing the EC method in a stellar evolution code. Our implementation in Cesam2k20 is tested with the solar model. Then we are able to model the α Cen binary system and compare the results with those from the original implementation of the method in the YREC code (Spada et al. 2018; Spada & Demarque 2019; Spada et al. 2021).

We use a physical description as close as possible to the one used in the YREC models and performed an optimization based on the Levenberg-Marquardt algorithm. The use of the entropy calibration ties α to the physics of near-surface convection as modeled by RHD simulations. This consequence has many advantages. First, having a value of α that changes with evolution is expected because there is no reason to think that the convection should keep the same properties as a star experiences strong structural modifications. Then, since with the entropy calibration the modeler does not have to provide a value for α , one can remove it from the set of adjustable parameters in an optimization process. This leads to one of two effects: either it just reduces the set of parameters and thereby facilitates the optimization or (instead of α) another free parameter can be adjusted in order to constrain another physical process taking place in a stellar model. The EC modeling not only leads to different results

than standard modeling, but also improves the physical description of convective envelope. However, it must be stressed that we do not expect any improvement of the seismic quality (e.g., a reduction of the surface effect) of stellar models by using EC models. Indeed, reliable stellar oscillations strongly depend on the modeling of the atmosphere which relies on the right choices for the opacities, the $T(\tau)$ relation and, the convective formalism. The EC models only bring stronger constraints on the free parameters of the latter.

Despite the numerical differences of the two codes, and a different choices of entropy prescriptions and corrections, CEsam2k20 was able to find optimal models of α Cen very similar to those obtained with YREC. Regarding this comparison, two important observations can be made: (1) the discrepancies between standard and EC models are larger for α Cen B, which is the less massive star in the system, and (2) the depths of the convective envelopes are slightly different between the standard and EC models. The first observation may be explained by the fact that the Hayashi track of α Cen B was computed outside the range of definition of the M13. In addition, a preliminary study (presented in a forthcoming paper) suggests that the entropy of the adiabat of M dwarfs stars are not fitted accurately by the prescriptions presented in this paper. This point emphasizes the fact that entropy prescriptions should not be used outside their range of definition. A study focused on PMS and M Dwarf stars will be the topic of a future work. The second observation may be of strong importance for the understanding of chemical and angular momentum transport because the location of the convective envelope boundary has a strong impact on it. This should also be a focus of future studies.

Acknowledgements. We are very grateful to Federico Spada for all the help he provided and for stimulating discussions. L.M. and L.G. acknowledge support from the Max Planck Society (MPG) under project “Preparations for PLATO Science” and from the German Aerospace Center (DLR) under project “PLATO Data Center”. L.M. acknowledges support from Agence Nationale de la Recherche (ANR) grant ANR-21-CE31-0018. A.S. acknowledges grants from the Spanish program Unidad de Excelencia María de Maeztu CEX2020-001058-M, 2021-SGR-1526 (Generalitat de Catalunya), and support from ChETEC-INFRA (EU project no. 101008324). J.K. acknowledges support from European Social Fund (project No. 09.3.3-LMT-K-712-19-0172) under grant agreement with the Research Council of Lithuania. Our study was partly supported by the European Union’s Horizon 2020 research and innovation program under grant agreement no. 101008324 (ChETEC-INFRA). H.-G.L. acknowledges financial support by the Deutsche Forschungsgemeinschaft (DFG, German Research Foundation) – Project-ID 138713538 – SFB 881 (“The Milky Way System”, subproject A04).

References

- Aikawa, M., Arai, K., Arnould, M., Takahashi, K., & Utsunomiya, H. 2006, in *Frontiers in Nuclear Structure, Astrophysics, and Reactions*, eds. S. V. Harissopulos, P. Demetriou, & R. Julin, *AIP Conf. Ser.*, **831**, 26
- Arnett, W. D., Meakin, C., Viallet, M., et al. 2015, *ApJ*, **809**, 30
- Asplund, M., Grevesse, N., Sauval, A. J., & Scott, P. 2009, *ARA&A*, **47**, 481
- Badnell, N. R., Bautista, M. A., Butler, K., et al. 2005, *MNRAS*, **360**, 458
- Ball, W. H. 2021, *Res. Notes Am. Astron. Soc.*, **5**, 7
- Böhm-Vitense, E. 1958, *Z. Astrophys.*, **46**, 108
- Broggini, C., Bemmerer, D., Caciolli, A., & Trezzi, D. 2018, *Prog. Part. Nucl. Phys.*, **98**, 55
- Canuto, V. M., & Mazzitelli, I. 1991, *ApJ*, **370**, 295
- Canuto, V. M., & Mazzitelli, I. 1992, *ApJ*, **389**, 724
- Canuto, V. M., Goldman, I., & Mazzitelli, I. 1996, *ApJ*, **473**, 550
- Castro, M., Baudin, F., Benomar, O., et al. 2021, *MNRAS*, **505**, 2151
- Christensen-Dalsgaard, J. 1982, *MNRAS*, **199**, 735
- Demarque, P., Guenther, D. B., Li, L. H., Mazumdar, A., & Straka, C. W. 2008, *Ap&SS*, **316**, 31
- Freytag, B., Steffen, M., & Dorch, B. 2002, *Astron. Nachr.*, **323**, 213
- Freytag, B., Steffen, M., Ludwig, H.-G., et al. 2012, *J. Comput. Phys.*, **231**, 919
- Gabriel, M., & Belkacem, K. 2018, *A&A*, **612**, A21
- Gough, D. O. 1977, *ApJ*, **214**, 196
- Gough, D. O., & Weiss, N. O. 1976, *MNRAS*, **176**, 589
- Grevesse, N., & Sauval, A. J. 1998, *Space Sci. Rev.*, **85**, 161
- Grigahcène, A., Dupret, M. A., Gabriel, M., Garrido, R., & Scuflaire, R. 2005, *A&A*, **434**, 1055
- Heney, L. G., Lelevier, R., & Levé, R. D. 1955, *PASP*, **67**, 154
- Heney, L., Vardya, M. S., & Bodenheimer, P. 1965, *ApJ*, **142**, 841
- Iglesias, C. A., & Rogers, F. J. 1996, *ApJ*, **464**, 943
- Ireland, L. G., & Browning, M. K. 2018, *ApJ*, **856**, 132
- Irwin, A. W. 2012, FreeEOS: Astrophysics Source Code Library [record ascl:1211.002]
- Kervella, P., Bigot, L., Gallenne, A., & Thévenin, F. 2017, *A&A*, **597**, A137
- Komm, R., Mattig, W., & Nesis, A. 1991, *A&A*, **252**, 812
- Krishna Swamy, K. S. 1966, *ApJ*, **145**, 174
- Kupka, F., & Muthsam, H. J. 2017, *Liv. Rev. Comput. Astrophys.*, **3**, 1
- Lebreton, Y., & Reese, D. R. 2020, *A&A*, **642**, A88
- Li, Y. 2012, *ApJ*, **756**, 37
- Ludwig, H.-G., Freytag, B., & Steffen, M. 1999, *A&A*, **346**, 111
- Ludwig, H.-G., Caffau, E., Steffen, M., et al. 2009, *Mem. Soc. Astron. Ital.*, **80**, 711
- Magic, Z., Collet, R., Asplund, M., et al. 2013, *A&A*, **557**, A26
- Magic, Z., Weiss, A., & Asplund, M. 2015, *A&A*, **573**, A89
- Mamajek, E. E., Prsa, A., Torres, G., et al. 2015, IAU 2015Resolution B3 on Recommended Nominal Conversion Constants for Selected Solar and Planetary Properties, https://www.iau.org/static/resolutions/IAU2015_English.pdf
- Marques, J. P., Goupil, M. J., Lebreton, Y., et al. 2013, *A&A*, **549**, A74
- Michaud, G. J., & Proffitt, C. R. 1993, in *IAU Colloq. 138: Peculiar versus Normal Phenomena in A-type and Related Stars*, eds. M. M. Dworetzky, F. Castelli, & R. Faraggiana, *ASP Conf. Ser.*, **44**, 439
- Mihalas, D., Dappen, W., & Hummer, D. G. 1988, *ApJ*, **331**, 815
- Morel, P. 1997, *A&AS*, **124**, 597
- Morel, P., & Lebreton, Y. 2008, *Astrophys. Space Sci.*, **316**, 61
- Morel, P., Provost, J., Lebreton, Y., Thévenin, F., & Berthomieu, G. 2000, *A&A*, **363**, 675
- Nordlund, A., & Galsgaard, K. 1995, *A 3D MHD code for Parallel Computers*, Tech. Rep., Astronomical Observatory, Copenhagen University, 1
- Nordlund, Å., & Stein, R. F. 1990, *Comput. Phys. Commun.*, **59**, 119
- Pasetto, S., Chiosi, C., Cropper, M., & Grebel, E. K. 2014, *MNRAS*, **445**, 3592
- Porto de Mello, G. F., Lyra, W., & Keller, G. R. 2008, *A&A*, **488**, 653
- Rogers, F. J., & Iglesias, C. A. 1992, *ApJ*, **401**, 361
- Rogers, F. J., & Nayfonov, A. 2002, *ApJ*, **576**, 1064
- Seaton, M. J. 2005, *MNRAS*, **362**, L1
- Seaton, M. J. 2007, *MNRAS*, **382**, 245
- Seaton, M. J., Yan, Y., Mihalas, D., & Pradhan, A. K. 1994, *MNRAS*, **266**, 805
- Serenelli, A. 2016, *Eur. Phys. J. A*, **52**, 78
- Sonoi, T., Ludwig, H. G., Dupret, M. A., et al. 2019, *A&A*, **621**, A84
- Spada, F., & Demarque, P. 2019, *MNRAS*, **489**, 4712
- Spada, F., Demarque, P., Basu, S., & Tanner, J. D. 2018, *ApJ*, **869**, 135
- Spada, F., Demarque, P., & Kupka, F. 2021, *MNRAS*, **504**, 3128
- Steffen, M. 1993, in *IAU Colloq. 137: Inside the Stars*, eds. W. W. Weiss, & A. Baglin, *ASP Conf. Ser.*, **40**, 300
- Tanner, J. D., Basu, S., & Demarque, P. 2016, *ApJ*, **822**, L17
- Thoul, A., Scuflaire, R., Noels, A., et al. 2003, *A&A*, **402**, 293
- Trampedach, R. 2010, *Ap&SS*, **328**, 213
- Trampedach, R., Stein, R. F., Christensen-Dalsgaard, J., Nordlund, Å., & Asplund, M. 2014, *MNRAS*, **445**, 4366
- Vernazza, J. E., Avrett, E. H., & Loeser, R. 1981, *ApJS*, **45**, 635
- Vinyoles, N., Serenelli, A. M., Villante, F. L., et al. 2017, *ApJ*, **835**, 202
- Wedemeyer, S., Freytag, B., Steffen, M., Ludwig, H. G., & Holweger, H. 2004, *A&A*, **414**, 1121
- Wolf, B. E. 1983, *A&A*, **127**, 93
- Zwintz, K., & Steindl, T. 2022, *Front. Astron. Space Sci.*, **9**, 914738

Appendix A: CIFIST grid for $[\text{Fe}/\text{H}] = 0.0$

Table A.1. Main characteristics of atmosphere models with solar metallicity in the CIFIST grid.

T_{eff} [K]	$\log g$ [cm s^{-2}]	s_{RHD} [$10^9 \text{ erg g}^{-1} \text{ K}^{-1}$]	s_{bot} [$10^9 \text{ erg g}^{-1} \text{ K}^{-1}$]	T_{bot} [K]	$\log_{10} p_{\text{bot}}$ [dyn cm^{-2}]	ρ_{bot} [$10^{-5} \text{ g cm}^{-3}$]
6234.05	4	2.167	2.16394	25382.0	7.20509	0.51801
4017.85	1.5	2.347	2.33927	17506.8	5.92127	0.04109
6485.56	4	2.377	2.37243	30518.6	6.97112	0.23571
6724.63	4.25	2.366	2.36573	100759	8.52705	2.44203
5926.77	4	1.992	1.99046	24617.6	7.53879	1.23210
5488.30	4.5	1.693	1.69257	24230.8	8.12577	5.66237
4968.68	2.5	2.3656	2.33936	15693.8	5.57400	0.02156
4924.10	3.5	1.827	1.82572	20075.0	7.33025	1.09334
6233.07	4.5	1.893	1.89205	26068.3	7.88415	2.64366
4774.62	3.2	1.878	1.86767	14170.1	6.05283	0.09360
5884.96	3.5	2.289	2.28402	24098.6	6.79680	0.20937
6456.55	4.5	1.992	1.99047	26268.5	7.68594	1.58967
6431.15	4.25	2.143	2.14145	29361.6	7.54681	0.96039
5226.43	4.25	1.704	1.70357	23381.4	8.00951	4.52138
4954.90	4	1.703	1.70253	22374.0	7.89403	3.68594
4980.91	4.5	1.613	1.61224	19202.3	7.66689	2.85028
4476.99	4	1.623	1.62227	18236.3	7.48812	2.01844
5537.34	4.43933	1.718	1.71374	15845.2	6.81084	0.49295
6084.98	4.43933	1.862	1.85267	15847.9	6.50204	0.22516
5475.94	4	1.827	1.82586	22047.9	7.59745	1.77328
5775.01	4.43933	1.7734	1.76747	15828.4	6.68128	0.35625
5774.02	4.43933	1.7807	1.77406	15702.4	6.63837	0.32537
6227.31	4	2.167	2.16403	25389.2	7.20546	0.51826
4476.44	2.5	2.042	2.03857	18263.8	6.63023	0.22395
3963.69	4.5	1.473	1.47209	13568.4	7.10341	1.32621
5866.50	4.5	1.777	1.77645	25857.8	8.10968	4.76272
4510.49	4.5	1.553	1.55222	16874.8	7.44482	2.11165
5432.51	3.5	2.022	2.01999	22513.4	7.25906	0.71921

Note: The complete table is available at CDS. The entropy s_{RHD} is the input entropy of the model, namely, the entropy of the inflow. Quantities with the subscript 'bot' are averages at the bottom of the simulation domain.

Appendix B: New coefficients for entropy functional representations

Table B.1. Coefficients for fit to CIFIST entropies at $[\text{Fe}/\text{H}] = 0.0$ to the functional form, Eq. (1) by Ludwig et al. (1999).

a_0	1.67199769
a_1	0.10100425
a_2	1.53800252
a_3	-1.41784705
a_4	0.10501243
a_5	-0.14985532

Table B.2. Coefficients for fit to CIFIST entropies to the functional form, Eq. (2) by Magic et al. (2013).

	$i = 0$	$i = 1$	$i = 2$
a_i	1.679013	0.051669	0.0080435
b_i	0.106697	-0.014053	0.0049438
c_i	-0.155069	-0.018509	-0.0028387
d_i	0.094921	0.039575	0.0065080
e_i	1.592803	-0.195275	-0.0015400
f_i	-1.460810	0.094490	0.0024214

Table B.3. Coefficients for fit to CIFIST entropies to the functional form, Eq. (4) by Tanner et al. (2016).

$[\text{Fe}/\text{H}]$	-3.0	-2.0	-1.0	0.0
A	0.989612	0.994939	1.004997	1.000814
B	-0.060934	-0.066989	-0.077271	-0.081859
s_0	1.346857	1.269913	1.398723	1.457433
x_0	3.560125	3.550494	3.518347	3.516223
β	1.196213	1.027772	0.731657	1.233787
τ_0	0.069621	0.103149	0.092566	0.086339

Table B.4. Values of mean molecular weight for fully ionized gas in the CIFIST grid.

$[\text{Fe}/\text{H}]$	μ_{RHD}
-4.0	0.593737
-3.5	0.593740
-3.0	0.593748
-2.5	0.593773
-2.0	0.593851
-1.5	0.594099
-1.0	0.594884
-0.5	0.596241
0.0	0.599408
0.5	0.611516

Note: These values are used for the computation of the f_μ factor of Eq. (9).