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LETTER TO THE EDITOR

Discovery of the first olivine-dominated A-type asteroid family

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ABSTRACT

The classical theory of differentiation states that due to the heat generated by the decay of radioactive elements, some asteroids form an iron core, an olivine-rich mantle, and a crust. The collisional breakup of these differentiated bodies is expected to lead to exposed mantle fragments, creating families of newly-formed asteroids. Among these new objects, some are expected to show an olivine-rich composition in spectroscopic observations. However, several years of spectrophotometric surveys have led to the conclusion that olivine-rich asteroids are rare in the asteroid main belt, and no significant concentration of olivine-rich bodies in any asteroid family has been detected to date. Using ESA's *Gaia* DR3 reflectance spectra, we show that the family (36256) 1999 XT17 presents a prominence of objects that are likely to present an olivine-rich composition (A-type spectroscopic class). If S-complex asteroids as the second most prominent spectroscopic class in the family are real family members, then arguably the 1999 XT17 family has originated from the break-up of a partially differentiated parent body. Alternatively, if the S-complex asteroids are interlopers, then the 1999 XT17 family could have originated from the breakup of an olivine-rich body. This body could have been part of the mantle of a differentiated planetesimal, which may have broken up in a different region of the Solar System, and one of its fragments (i.e. the parent body of the 1999 XT17 family) could have been dynamically implanted in the main belt.

Key words. techniques: spectroscopic – catalogs – minor planets, asteroids: general

1. Introduction

Early-generation planetesimals are bodies that formed $\lesssim 2$ Myr after the formation of our Solar System. They are believed to have gone through a process of differentiation (Šrámek et al. 2012; Kruijer et al. 2014), which causes a body to organise into layers of different densities and composition, due to the heat generated by the decay of radioactive elements, such as ^{26}Al . In the classical scenario of asteroid differentiation, these layers comprise a basaltic crust, an olivine-rich mantle, and an iron core (Elkins-Tanton & Weiss 2017).

Asteroid differentiation is corroborated by the existence of meteorites originating from the different aforementioned layers, such as iron meteorites (Cloutis et al. 1990), mesosiderites (Hewins 1983; Greenwood et al. 2006), brachinites and pallasites (Cruikshank & Hartmann 1984), and andesites (Barrat et al. 2021), as well as the howardite-eucrite-diogenite (HED) meteorites (McCord et al. 1970). However, apart from few clear cases such as the asteroid (4) Vesta (Russell et al. 2012), the detection of differentiated asteroids and the assessment of their abundance in the main belt remains a challenge.

One way to detect a potentially differentiated asteroid is to discover a very-high-density object, indicating the presence of an iron core. Although large improvements were made in the last two decades in the measurement of asteroid sizes¹ and volumes (e.g. Durech et al. 2010; Āurech et al. 2011; Vernazza et al. 2021), mass estimations (Fienga et al. 2020) are still limited and

may suffer from large uncertainties. This issue leads to a reliable density estimation for a relatively small number of objects in the main belt and very-high-density objects detected may be due to errors (Carry 2012). A related issue is that large collisions may have induced voids in a body, causing a potential high density to be camouflaged and thus hiding hints of differentiation.

Another way to discover differentiated objects is to study collisional families of asteroids. In principle, if a large impact event shatters a differentiated asteroid, the resulting family of newly born smaller asteroids should contain fragments sampling the layers of the parent body. Hence, a variety of compositions should be detectable within the collisional family. However, a catastrophic disruption of the parent body is needed to reveal such layers. Indeed, cratering events resulting from impacts may only produce ejecta from the parent body's crust. One example is the Vesta family, which resulted from the two impacts that formed the large basins Rheasilvia and Veneneia on (4) Vesta (Marchi et al. 2012). Remote sensing data from NASA's Dawn space mission (Russell et al. 2012), ground-based spectroscopy, and geometric visible albedo data, have revealed a remarkable link between the crust of (4) Vesta, its family members, and the HED meteorites. The basaltic composition of Vesta family members indicates that Vesta is a differentiated asteroid, but, despite its large extension, Vesta family members do not clearly show mantle and core material compositions.

Spectroscopic searches for traces of differentiated material in the main belt have led to the discovery of basaltic (McCord et al. 1970), olivine-rich (DeMeo et al. 2019), and potentially metal-rich asteroids (Fornasier et al. 2010; Harris & Drube 2014). However, different spectroscopic surveys shed light on the

¹ See references at mp3c.oca.eu.

scarcity of olivine-rich asteroids compared to other asteroid types in the main belt. This lack of mantle-like asteroids is known as the ‘missing mantle problem’ (Burbine et al. 1996). Olivine-rich asteroids show a moderately high albedo, and a spectrum characterised by extremely reddish slopes shortwards of $0.7\,\mu\text{m}$, a strong absorption feature centred at about $1.05\,\mu\text{m}$, and an absent or weak absorption feature at $2\,\mu\text{m}$. These objects are classified as A-types in most taxonomic schemes (Bus & Binzel 2002; DeMeo et al. 2009; Mahlke et al. 2022).

Using spectroscopic observations in the near-infrared (NIR) wavelength range, DeMeo et al. (2019) found that the olivine-rich asteroid abundance in the main belt is as low as 0.16%, for asteroids with diameters $D > 2\,\text{km}$. In addition, olivine-rich asteroids are found to be evenly distributed throughout the main belt, showing no statistically significant concentration in any asteroid family (see Sanchez et al. 2014; DeMeo et al. 2019). Several theories have been developed to explain the aforementioned observations: i) A-type asteroids are highly sensitive to space-weathering phenomena, and they can no longer be detected (Popescu et al. 2018b); ii) olivine-rich objects were shattered into smithereens by collisions and they have become too small to be detected by current observational techniques (Sanchez et al. 2014; DeMeo et al. 2019); iii) differentiated planetesimals formed early and accreted into planets in their majority, leaving behind very few differentiated bodies to be implanted in the main belt as intact bodies or fragments; iv) the olivine is masked by some material of another type on the asteroids (DeMeo et al. 2019); or v) current theories about planetesimal differentiation and formation of thick olivine mantles are incorrect (DeMeo et al. 2019).

In this work, we use the 60 518 Solar System objects (SSO) reflectance spectra included in the *Gaia* Data Release 3 (DR3) (Gaia Collaboration 2023) to study the abundance of olivine-rich asteroids in families. In particular, by using the *Gaia* DR3 and ground-based asteroid spectra, we test the hypothesis that there is no significant concentration of A-type asteroids in any asteroid family in the main belt. This work is focused primarily on the catalogue of asteroid family membership from Nesvorný et al. (2015). In the following, we present the dataset used to select the family of asteroids most likely to contain a high abundance of A-types. Then, we characterise this family using the *Gaia* DR3, literature data, and a newly acquired reflectance spectrum. We discuss our findings in the final section.

2. Family selection

Using the families and family memberships from Nesvorný et al. (2015), and the asteroid classification from several published sources (see Appendix A), we calculated the proportion of potential A-type asteroids in each asteroid family. To do so, we calculated the ratio between the number of objects classified at least once as A-type in different taxonomies, or as Ad-type from Popescu et al. (2018a), and the total number of asteroids with any other spectral type in each family. Then, we considered only the families having a potential A-type abundance of above 10%. These families are (10811) Lau (FIN 619), (36256) 1999 XT17 (FIN 629), and (7468) Anfimov (FIN 635), which were found to contain 13, 23 and 28% of potential A-type members, respectively. We looked into the *Gaia* DR3 data set (Gaia Collaboration 2023) for spectra of these families’ members. We found that these families have 1, 15, and 7 asteroids with a spectrum in the DR3 dataset, respectively. Since only asteroid (10811) Lau itself shows a spectrum in this data set, we did not consider the Lau family in this study. For the two

other families, we studied the average signal-to-noise ratio (S/N) calculated by (Gaia Collaboration 2023), of the DR3 spectra of the members of each family. We found that the Anfimov family shows a mean S/N of 23, with asteroids (7468) Anfimov and (18853) 1999 RO92 having the spectra with the maximum S/N, respectively of 47.41 and 32.21. These two objects had already been characterised based on spectroscopy (DeMeo et al. 2019; Mahlke et al. 2022), with the other members showing a low S/N; thus, we chose to focus on the family (36256) 1999 XT17 for the purposes of this study. This family shows an average S/N of 40, with nine objects having a $S/N > 21$, which is considered a threshold for reliable spectral classification in the visible (VIS) wavelength range (Delbo et al. 2012). This family appears to be the most suited for our search for potential A-type members. The 1999 XT17 family is located in the so-called ‘pristine zone’ of the main belt (Brož et al. 2013). This zone extends between 2.825 and 2.955 au and is bounded by the 5:2 and 7:3 mean motion resonances (MMRs) with Jupiter. The 1999 XT17 family contains 58 members and it has been reported to be a S-complex family by Nesvorný et al. (2015).

3. Characterisation of the (36256) 1999 XT17 family

We searched in the literature for additional spectral data for the 15 members of the 1999 XT17 family having a DR3 reflectance spectrum (Table 1). We found a NIR spectrum for asteroid 36256, from observations of DeMeo et al. (2019). For asteroids 25356 and 40671, we used NIR observations from the MOVIS catalogue (Popescu et al. 2018a); and for asteroids 36256 and 99004, we found visible-light spectrophotometry data from the Sloan Digital Sky Survey (SDSS; DeMeo & Carry 2013). We also performed NIR spectroscopic observations at the NASA Infrared Telescope Facility (IRTF) of the asteroid 33763 (Appendix B).

Using the aforementioned spectra, we followed the method of Avdellidou et al. (2022) to classify each 1999 XT17 family member in the Bus-DeMeo taxonomy (DeMeo et al. 2009), as this scheme extends to the NIR. For each asteroid, a reflectance curve-matching was performed between a given spectrum and the spectral class templates of DeMeo et al. (2009), by minimising the value of a reduced χ^2 figure of merit. For example, asteroid 36256 has a VIS *Gaia* DR3 and a NIR spectrum, as well as SDSS spectrophotometry. Therefore, three independent χ^2 were calculated for this object, between the three independent data sets and the Bus-DeMeo templates. The DR3 spectra were considered from 0.45 to $0.90\,\mu\text{m}$ only. Then, the reduced χ^2 obtained for each data set were summed, to obtain a global value of this figure of merit. The two classes with the lowest value of the global χ^2 represent the best two classes for a given asteroid (details in Appendix B of Avdellidou et al. 2022). We visually inspected each spectrum to check the best matching classes given by the automatic procedure. The best two classes obtained are presented in Table 1. Our algorithm classified 12 asteroids out of 15 to be A-types, as first or second best match. The remaining three objects, namely, asteroids 34902, 83124, and 88057, were classified as S-complex objects as the first or second best match.

In the literature, 7 objects out of the 15 that have a spectrum in the DR3 dataset have been classified. Among them, only asteroid 36256 has been characterised by DeMeo et al. (2019) using NIR spectroscopy, and is found A-type by DeMeo et al. (2019) and Mahlke et al. (2022) in their respective taxonomic schemes. Asteroids 25356 and 40671 are found to be Ad types by Popescu et al. (2018a), from MOVIS NIR colours;

Table 1. Members of the (36256) 1999 XT17 family with a *Gaia* DR3 spectrum.

Asteroid Number	<i>D</i> (km)	<i>p_V</i>	Avg <i>S/N</i>	<i>H</i>	Spectral class
15610	5.84	0.30	35.32	13.3	A, L
16789	6.62	0.18	38.25	13.5	A, Sv
20975	5.24	0.28	36.01	13.6	A, Sv
25356	–	–	66.49	12.8	A, Sa
27565	5.92	0.19	20.74	13.6	A, L
33763	–	–	48.07	13.4	A, Sa
34902	4.15	0.15	14.7	14.4	Sq, Sr
36256	10.21	0.19	66.56	12.4	A, Sa
40671	–	–	28.09	13.6	A, Sa
57276	7.20	0.31	128.7	12.9	A, Sv
66676	6.55	0.22	52.28	13.3	L, A
83124	–	–	14.88	14.7	S, Sr
88057	–	–	20.75	14.5	K, S
99004	3.83	0.10	13.93	14.6	L, A
140349	3.69	0.20	13.86	14.7	Sv, A

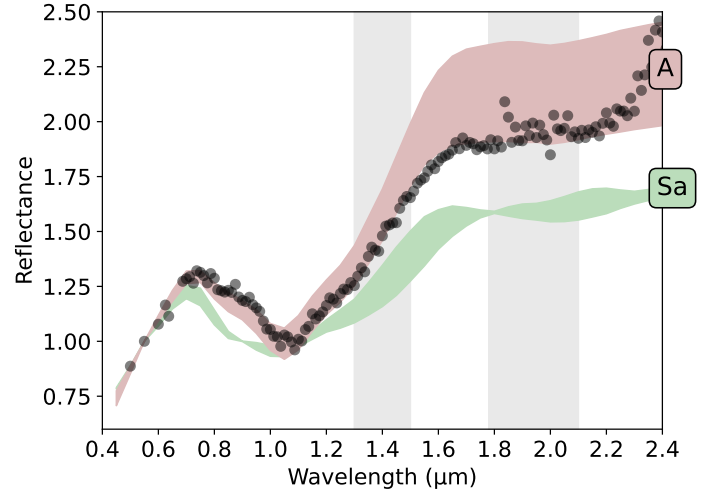
Notes. Here is given the diameter *D* of family members, their geometric visible albedo *p_V*, the average signal-to-noise ratio of their DR3 spectrum *S/N*, their *H* magnitude, and the two best spectral classes given by our classification algorithm. The physical properties data were obtained from mp3c.oca.eu.

and asteroids 27565, 34902, 66676, and 99004 are found to be part of the S-complex from spectrophotometric observations (Carvano et al. 2010; DeMeo & Carry 2013; Sergeyev & Carry 2021; Sergeyev et al. 2022). Figure E.1 shows the *Gaia* DR3 spectra of the 15 family members, along with supplementary data when available, including the NIR spectrum of 33763 that we obtained. The latter confirms that 33763 is an A-type.

Among the 43 members of the 1999 XT17 family that do not have a DR3 spectrum, 21 are classified in the literature based on spectrophotometric data in the VIS or NIR wavelength range. Among these, asteroid 76627 is classified as Ad from NIR colours (Popescu et al. 2018a), while 58777 and 201232 are classified as A (Carvano et al. 2010; Sergeyev et al. 2022). Asteroid 254896 is found A by Carvano et al. (2010) and S by Sergeyev et al. (2022). Seven asteroids (40427, 204222, 222898, 223304, 271103, 285502, and 362895) are classified as S-complex members (Carvano et al. 2010; DeMeo & Carry 2013; Sergeyev & Carry 2021; Sergeyev et al. 2022), and asteroid 62676 is classified as S-type by Sergeyev et al. (2022) and X-type by Carvano et al. (2010). The remaining objects are classified as L, K, C, D, V or X-types (not considering the objects classified as U here).

In total, 36 members of the 1999 XT17 family have a classification. Among them, 16 are classified as A-type, which makes a proportion of potential A-types of 44.4%. Among these, four asteroids are confirmed A-types from the combination of DR3 and NIR data. The second most abundant class in this family is the S-complex, with a proportion of potential S-complex asteroids of 30.5%.

For each asteroid with a *Gaia* DR3 spectrum, we combined the independent data sets used in the classification to produce a single spectrum per object. Then, following the procedure of Avdellidou et al. (2022), we averaged the single spectra of the 12 asteroids that we classified as A-types, to produce a family averaged reflectance spectrum (Fig. 1) that shows an excellent match to the A-type template of the Bus-DeMeo taxonomy.

**Fig. 1.** Average spectrum of the (36256) 1999 XT17 family members classified here as A-types, along with the A (red area) and Sa (green area) Bus-DeMeo templates.

Using literature data², the uncertainty-weighted average of the geometric visible albedo (*p_V*) of the potential A-type members of the family was calculated to be 0.22, with an rms value of 0.07. This value is within the range of the A-types' *p_V* (Mahlke et al. 2022).

To gain insights into the composition of these family members, we compared the two VNIR spectra of asteroids 36256 and 33763 with those of the meteorites and samples available in the Relab database (see Appendix C for details). As the best spectral analogues, we obtained mostly olivine assemblages and olivine-rich meteorites for both asteroids, confirming previous findings on other A-types (DeMeo et al. 2022). Calculating the equivalent geometric visible albedo (Beck et al. 2021) of the Relab meteorite samples and comparing it with the family's mean *p_V*, following the method developed in Avdellidou et al. (2022), we found that the CK6 meteorite LEW87009, the R-chondrites Rumuruti and LAP04840, the R4 chondrite MIL07440, the olivine of the shergottite ALHA77005, and the brachinites NWA4882 and ALH84025, are the best matches for asteroids 36256 and 33763. However, the shergottite ALHA77005 is identified as a martian meteorite; thus, we conclude it does not come from the outer-main belt 1999 XT17 family.

To gain further information on the characterised family members, we studied their positions in the proper orbital element space, and in the *H* versus proper semimajor axis *a* space. In the same spaces, we display the neighbouring (221) Eos family as defined by Nesvorný et al. (2015). In the *H* versus *a* space (panel c of Fig. E.2), the 1999 XT17 family members with *H* ≤ 14 lay outside the V-shape of the Eos family. Interestingly, we classified them as A-types and they have a *S/N* > 21. In the *e* versus *a* space (panel a of Fig. E.2), we see that the members of the 1999 XT17 family have a global proper eccentricity higher than the majority of Eos family members. In the sin *i* versus *a* space (panel b of Fig. E.2), however, members of the 1999 XT17 family show similar proper inclination values to the Eos family members.

² mp3c.oca.eu.

4. Discussion

First, we investigated the possibility that our A-type classification based on *Gaia* data is affected by an issue in the DR3, which tends to overestimate the reflectance values mostly in the reddest bands (Galinier et al. 2023; Gaia Collaboration 2023). Due to this reddening, a non-A-type asteroid could potentially be classified as such, since this class shows a red reflectance in the VIS wavelength range (see Table D.1 for the confusion matrix on the A-type classification). We calculated a false positive rate (asteroids that are not A-type in the literature, but that are classified A-type by our method) smaller than 2.5%. Therefore, we concluded that it is unlikely that our classification is affected by the aforementioned DR3 reddening. Moreover, the true positive rate (asteroids that are A-types in the literature correctly classified here) is higher than 71% (Appendix D). This adds confidence to our classification methods for A-types. However, we suggest using spectral matching between the DR3 and Bus-DeMeo templates with caution overall, since we investigated only the A-type case here.

In addition, we found that four asteroids among the six false negatives (asteroids that are A-type in the literature not classified as A here) in the A-type confusion matrix (Appendix D) have a $S/N < 21$. This is below the threshold commonly accepted for a reliable classification (Delbo' et al. 2012). Thus, it is not impossible that the three objects of the 1999 XT17 family that we did not classify as A-types could, in fact, turn out to be real A-types. Similarly, the three objects with a low S/N that we classified as A-types have to be taken with more caution.

Assuming that our classification and that the literature classes assigned from spectrophotometry are correct, then their combination indicates that the dynamical 1999 XT17 family is mainly composed of A-type objects, with a secondary S-complex class. In the following, we explore the possibility that some of these S-complex objects are interlopers of the 1999 XT17 family.

First, we studied the S-complex families in the pristine zone, assuming that they could be the source of this secondary compositional class. According to Nesvorný et al. (2015), there are other five families in the pristine zone that show a S-complex composition: (158) Koronis, (832) Seraphina, (918) Itha, (10811) Lau, and (15477) 1999 CG1. However, none of these families is in close vicinity to the 1999 XT17 family, making them an unlikely origin for the contamination.

While the 1999 XT17 family resides in a relatively empty zone compared to other regions of the main belt, it was shown that this 'pristine zone' is contaminated by the Eos family (Tsirvoulis et al. 2018; Brož & Morbidelli 2019). The 1999 XT17 family is located just inwards of the 7:3 MMR with Jupiter, in the same region as Eos family members that drifted away from the centre of the family by the action of non-gravitational forces (Nesvorný et al. 2015). These 'Eos family fugitives' crossed the 7:3 MMR fast enough so that their eccentricity did not reach planet crossing values. This resonance crossing can lead to a discontinuity in eccentricity between Eos fugitives and the Eos core family members (Nesvorný et al. 2015). On the other hand, this resonance crossing affects little the proper inclination of the objects.

Interestingly, every 1999 XT17 family member belonging to the secondary S-complex lays inside the Eos family V-shape, and presents similar inclination values and a discontinuity in eccentricity compared to the Eos family core members (see panels a and b of Fig. E.2, and Brož & Morbidelli 2019). These objects could therefore be Eos family fugitives that were merged with

asteroids of the 1999 XT17 family by the hierarchical clustering method used for family identification (Nesvorný et al. 2015). Spectroscopically, Eos is mostly composed of K-type asteroids (Brož & Morbidelli 2019). However, the distinction between S and K-types is non trivial in the visible for the Eos family members (Vokrouhlický et al. 2006). We conclude that it is not impossible that the S-complex asteroids of the 1999 XT17 family are part of the Eos fugitives.

On the other hand, the cluster of asteroids laying outside the Eos family V-shape contains ten asteroids, nine of which are classified as A-types from their DR3 spectra with a $S/N > 21$. Asteroids 36256 and 33763 are confirmed A-types from NIR spectroscopy, while 25356 and 40671 are confirmed A-types from the combination of their *Gaia* DR3 spectrum and their NIR colours (Popescu et al. 2018a). Moreover, the link we established between asteroids 36256 and 33763 and olivine-rich samples and meteorites from the Relab database, confirms the likely high content in olivine of these bodies. Since these objects are out of the V-shape of Eos, they are less likely to be part of Eos family fugitives. We therefore argue that the A-type members of the 1999 XT17 family originate from the breakup of a common parent body.

If the true 1999 XT17 family is composed only by A-types, it could originate from the breakup of an olivine-rich asteroid that was once itself part of a differentiated planetesimal. The latter could have been catastrophically disrupted in another region and a fragment of it could subsequently have been implanted in its current location via a dynamical process (Bottke et al. 2006; Raymond & Izidoro 2017). Corroborating evidence could be the small size (10 km) of asteroid (36256) 1999 XT17, the largest member of the small-sized 1999 XT17 family (some tens of asteroids), coupled with the small diameters of the rest of its members (~5 km). This was recently shown to be a valid scenario for the parent body of the enstatite meteorites of subtype EL. It has been proposed that this family was delivered to the inner main belt from the terrestrial region during the early evolution phases of our Solar System (Avdellidou et al. 2022). On the other hand, some olivine-rich main belt asteroids could be the result of direct accretion of grains from an oxidized region of the solar nebula, instead of magmatic differentiation, as understood from the study of R-chondrites (Schulze et al. 1994). The olivine-rich parent body at the origin of the 1999 XT17 family could have formed through such processes.

We go on to explore the hypothesis that the 1999 XT17 family is totally or partially differentiated. As we state above, the second most abundant class in this family according to the literature is the S-complex. S-complex asteroids are typically associated to ordinary chondrite meteorites (Vernazza et al. 2015; Reddy et al. 2015; DeMeo et al. 2022), that are undifferentiated (see Vernazza et al. 2015, and references therein). Based on the hypothesis where these S-complex objects are correctly classified and belong to the 1999 XT17 family, it is possible that the parent body of the 1999 XT17 family was only partially differentiated. It could have built an olivine mantle without heating enough for the crust to melt, which may have preserved its undifferentiated nature (Weiss & Elkins-Tanton 2013). Unfortunately, spectrophotometry is often not precise enough to distinguish between spectroscopic classes. Further spectroscopic observations in VIS and NIR could help in distinguishing between our hypotheses, for example, following the methods of Sunshine et al. (2007).

In any case, spectroscopic studies that attempted to map the olivine-rich asteroids in the main belt concluded that there is no prominent concentration of olivine-rich objects in any asteroid

family. They are instead found to be spread throughout the main belt. Here, we identified the first concentration of A-type objects in an asteroid family, using *Gaia* DR3 data. This identification makes the 1999 XT17 family the first olivine-rich family in the main asteroid belt.

5. Conclusions

We report on the discovery of the first concentration of A-type asteroids in an asteroid family in the pristine zone of the outer main belt. The (36256) 1999 XT17 family shows the presence of 16 potential A-type asteroids, among which four have been confirmed by NIR data. The comparison with meteorites of asteroids 36256 and 33763 confirms that these objects should be mostly composed of pure olivine. Depending on the membership in the family of a few S-complex asteroids, we propose two main hypothesis for the origin of the family: (i) the parent body could have been a pure olivine-rich object from the mantle of a partially or totally differentiated planetesimal, which got disrupted possibly at a different location in the Solar System, before getting implanted in the pristine zone; or (ii) the family comes directly from a partially differentiated object that broke in situ. New NIR spectroscopic observations of family members could help disentangle our different scenarios.

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Appendix A: Spectral classes from the literature

For this work, we used asteroid spectral classes derived from spectroscopic observations from the following references: Tatsumi et al. (2022), Migliorini et al. (2021), Arredondo et al. (2021), Gartrelle et al. (2021), Hasegawa et al. (2021), De Prá et al. (2020a,b), Lucas et al. (2019), DeMeo et al. (2019), Devogèle et al. (2019), Morate et al. (2019), Popescu et al. (2019), Perna et al. (2018), Migliorini et al. (2018), De Prá et al. (2018), Devogèle et al. (2018), Morate et al. (2018), Borisov et al. (2017), Migliorini et al. (2017), Leith et al. (2017), Lucas et al. (2017), de León et al. (2016), Morate et al. (2016), Ribeiro et al. (2014), DeMeo et al. (2014), Thomas et al. (2014), Popescu et al. (2014), Oszkiewicz et al. (2014), Sanchez et al. (2013), Gietzen et al. (2012), de Sanctis et al. (2011), de León et al. (2010), DeMeo et al. (2009), Clark et al. (2009), Mothé-Diniz et al. (2008), Roig et al. (2008), Licandro et al. (2008), Mothé-Diniz & Nesvorný (2008), Alvarez-Candal et al. (2006), Lazzarin et al. (2005), Marchi et al. (2005), Mothé-Diniz et al. (2005), Duffard et al. (2004), Binzel et al. (2004), Lazzarin et al. (2004), Bus & Binzel (2002), Carvano et al. (2001). We also used asteroid spectral classes derived from spectrophotometric observations from the following references: Carvano et al. (2010), Popescu et al. (2018a), Sergeyev & Carry (2021), Sergeyev et al. (2022), DeMeo & Carry (2013)

Appendix B: NIR observation of (33763) 1999 RB84

On June 13 2023, we used the SpeX (Rayner et al. 2003) and MORIS instruments (Gulbis et al. 2011) at NASA's Infrared Telescope Facility (IRTF) to observe asteroid 33763 at 9.4° phase angle. SpeX was used with a slit 0.8×15" in PRISM mode, covering wavelengths between 0.7 and 2.5 μm with a spectral resolution of ~200 in a single configuration. The asteroid's magnitude was 17.4 and the exposure time for a single frame was 180 sec. We acquired 6 cycles 'AB-BA' frame pairs, for a total exposure time of 18 min, where the asteroid was shifted along the slit. Following well established procedures (Reddy et al. 2009), we calculated the asteroid's reflectance $R(\lambda)$ as a function of the wavelength λ using Eq. B.1:

$$R(\lambda) = \frac{A(\lambda)}{S_L(\lambda)} \times Poly\left(\frac{S_L(\lambda)}{S_T(\lambda)}\right), \quad (\text{B.1})$$

where $A(\lambda)$, $S_L(\lambda)$, $S_T(\lambda)$ are the wavelength-calibrated raw spectra respectively of the asteroid, of the local HD137782 G2V star observed within ~300" of the asteroid, and of the trusted solar analogue star SA107684 that was observed at a similar air-mass, respectively. The function $Poly()$ represents a polynomial fit of the stars ratio, excluding the regions affected by the telluric water vapour absorption ($1.3 < \lambda < 1.5$, $1.78 < \lambda < 2.1$, and $\lambda > 2.4 \mu\text{m}$). Asteroid spectra were shifted to sub-pixel accuracy to align with the calibration star spectra. In general, the local star ensures accurate removal of the telluric features, but may require the slope correction $Poly\left(\frac{S_L(\lambda)}{S_T(\lambda)}\right)$ due to difference between the local star's spectrum and the one of the Sun.

Appendix C: Comparison with meteorites

We used the full VNIR spectra of asteroids 36256 and 33763 to compare these objects with the samples and meteorites available in the Relab database to gain information on their composition. To do so, we used the exponential space-weathering model

derived by Brunetto et al. (2006) to correct the asteroids spectra from the effects of space weathering (as done for example in DeMeo et al. 2022; Mahlke et al. 2023). Brunetto et al. (2006) found that the ratio between irradiated and non-irradiated silicate samples shows almost no contribution from the spectral bands, and gives a continuum that can be fitted by an exponential curve:

$$W(\lambda) = K \exp\left(\frac{C_s}{\lambda}\right), \quad (\text{C.1})$$

where λ is the wavelength, K is a normalising scale factor that depends on the normalisation of the spectra, and C_s is the strength of the exponential curve, representing the strength of the space weathering. For each of the two asteroids, we calculated the ratio between their VNIR spectrum and the spectrum of every sample of the Relab database. To do so, we re-sampled the asteroids and the meteorites spectra from 450 to 2500 nm for asteroid 36256, and from 450 to 1800 nm for asteroid 33763. We adopted a narrower wavelength range for this latter asteroid, to filter out the noisier part of its NIR spectrum affected by a positive slope. Then, we fitted the exponential space-weathering model of Brunetto et al. (2006) to each calculated ratio, to compute de-weathered asteroids spectra.

We used a curve matching method to find the best analogues of asteroids 36256 and 33763, following established methods (see e.g. Avdellidou et al. 2022; DeMeo et al. 2022; Mahlke et al. 2023). We calculated the Φ_{comb} parameter of Popescu et al. (2012) between the Relab samples and the de-weathered asteroids spectra, and as the best matches, we obtained mostly olivine mineral spectra, similarly to the results from DeMeo et al. (2022). Among the meteorites found as best spectral matches, we found the CK6 meteorite LEW 87009, the R-chondrite Rumuruti, the shergottite ALHA77005, the pallasites Esquel and Thiel Mountains, and the brachinite EET99402 for asteroids 36256 and 33763.

We used the equivalent geometric visible albedo presented in Beck et al. (2021) to try to discriminate between these best spectral matches. We followed the method as implemented in Avdellidou et al. (2022): we determined the equivalent geometric visible albedo, p_V , of the Relab meteorite samples, to be able to compare their reflectance values to the mean albedo of the family, $p_V = 0.22$. We calculated the reflectance value at 0.55 μm of the meteorites observed at a phase angle of 30°, and we derived their reflectance value at a phase angle of 0°, using Eq. 2 of Beck et al. (2021). The comparison between this equivalent geometric visible albedo to the family's mean p_V gives us an indication of which meteorite samples might be the best matches for these family members. Among the best meteorite analogues given by the curve matching method, we found that the CK6 meteorite LEW87009, the R-chondrites Rumuruti and LAP04840, the R4 chondrite MIL07440, the olivine of the shergottite ALHA77005, and the brachinites NWA4882 and ALH84025, show an equivalent geometric albedo close to $p_V = 0.22$. These meteorites are therefore the best matches found of asteroids 36256 and 33763. However, the shergottite ALHA77005 is a martian meteorite, and is therefore not a reasonable match for this outer-main belt asteroid family.

Appendix D: Classification algorithm

We studied the possibility that the A-type classification driven by Gaia data got contaminated by objects affected by a Gaia DR3 reddening (Gaia Collaboration 2023; Galinier et al. 2023). To investigate this issue, we tested our algorithm by classifying

Table D.1. Confusion matrix of the classification of A-type asteroids by our algorithm. The list of 389 asteroids both classified by DeMeo et al. (2009, 2014, 2019) and having a DR3 spectrum considered for this study contains 21 A-type asteroids.

Predicted True	A-type	Non-A-type
A-type	15	8
Non-A-type	6	360

asteroids both having a spectrum in the DR3, and that have been classified from VIS-NIR observations by DeMeo et al. (2009, 2014, 2019). This list contains 389 asteroids. We applied our classification method on their DR3 spectra, in the wavelength range from 0.45 to 0.90 μm . This wavelength range was chosen in order to not take into account the first and last two bands of the DR3 spectra often affected by systematic errors Gaia Collaboration (2023), Oszkiewicz et al. (2023), and to limit the effect of a potential DR3 reddening. Moreover, a quality flag going from 0 (good quality) to 2 (poorer quality) was assigned to each of the 16 bands of the Gaia DR3 spectra. We did not consider the bands flagged with a non-zero number, as done in Galinier et al. (2023).

We studied the performances of our algorithm for the classification of A-types using the confusion matrix presented in Table

D.1. In the list of 389 asteroids, 21 are classified as A-types. We found that eight asteroids that are not A-type in DeMeo et al. (2009, 2014, 2019) are classified A-type by our method. This corresponds to a false positive rate of around 2.2%. However, among these false positives, three asteroids were later classified as A-type by Mahlke et al. (2022) in his own taxonomic scheme. If we consider this information, the false positive rate goes down to less than 1.5%.

Among the 21 A-type asteroids of the list, 15 get correctly classified as A-types as first or second best match with the Bus-DeMeo templates. This corresponds to a true positive rate of above 71%. The accuracy of the classification corresponds to the sum of the true positive and the true negatives, divided by the total population. The accuracy for the A-type classification here is therefore of $\frac{15+360}{389} = 96\%$.

Among the six A-type asteroids of the literature that did not get classified as A-type here (false negative rate of 28.6%), four asteroids have a $S/N < 21$. We conclude that the classification of low S/N A-type asteroids should be taken with more caution: it is not impossible that a low S/N asteroid that was not found A by our method is an A. Similarly, asteroids found A by our method and that show a $S/N < 21$ might not be real A.

We did not aim to extensively study the performance of our classification algorithm on other taxonomic types. We therefore advise caution with the use of spectral matching between DR3 spectra and Bus-DeMeo templates to classify asteroids in the DR3 dataset.

Appendix E: Figures

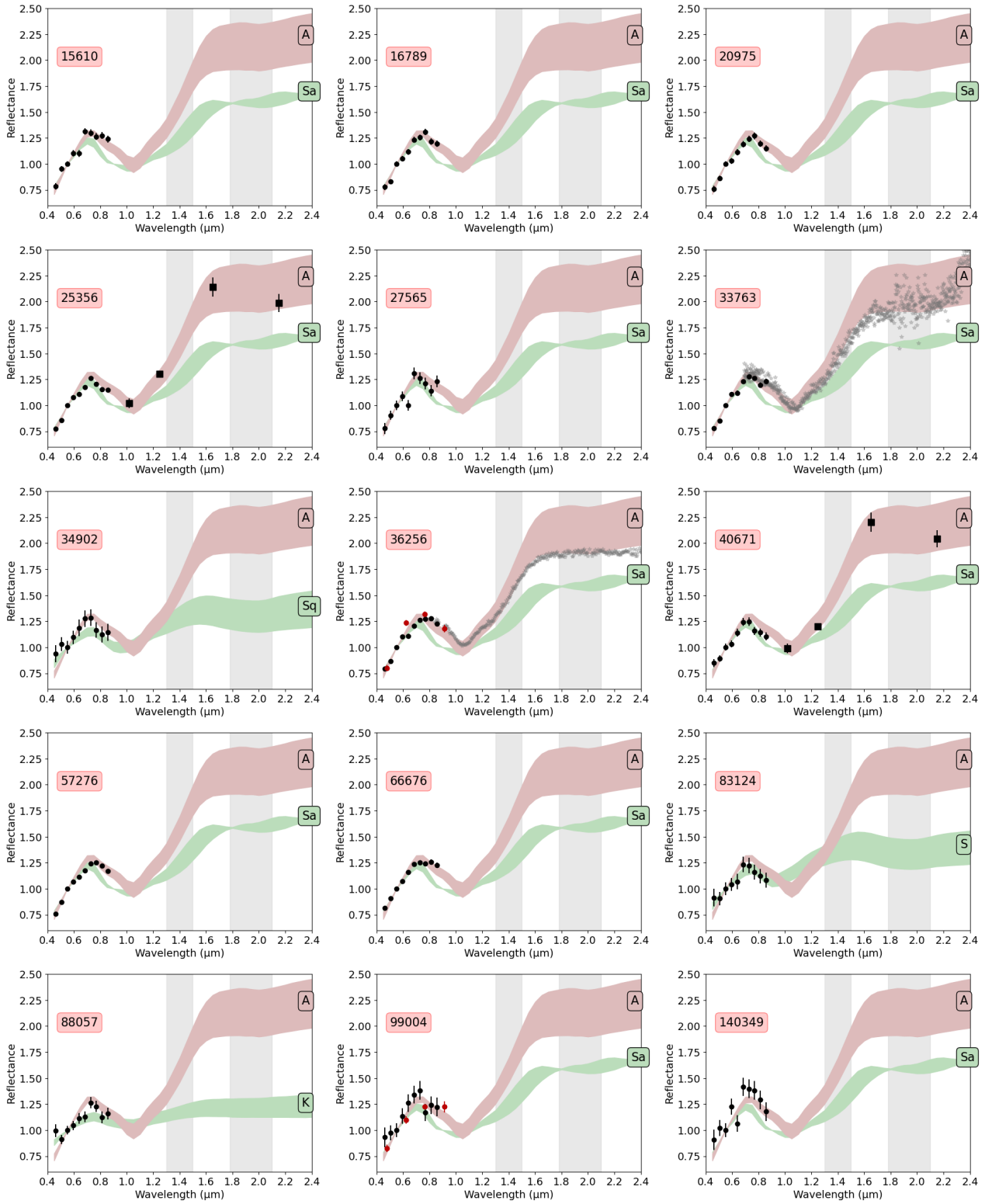


Fig. E.1. Available data of the members of the family (36256) 1999 XT17 in the Gaia DR3: Gaia spectra (black dots), SDSS data from [DeMeo & Carry \(2013\)](#) (red dots), MOVIS data from [Popescu et al. \(2018a\)](#) (black squares), and NIR spectra from [DeMeo et al. \(2019\)](#) (grey stars). The shaded coloured areas represent the reflectance spectra from the A-type and other classes templates of [DeMeo et al. \(2009\)](#).

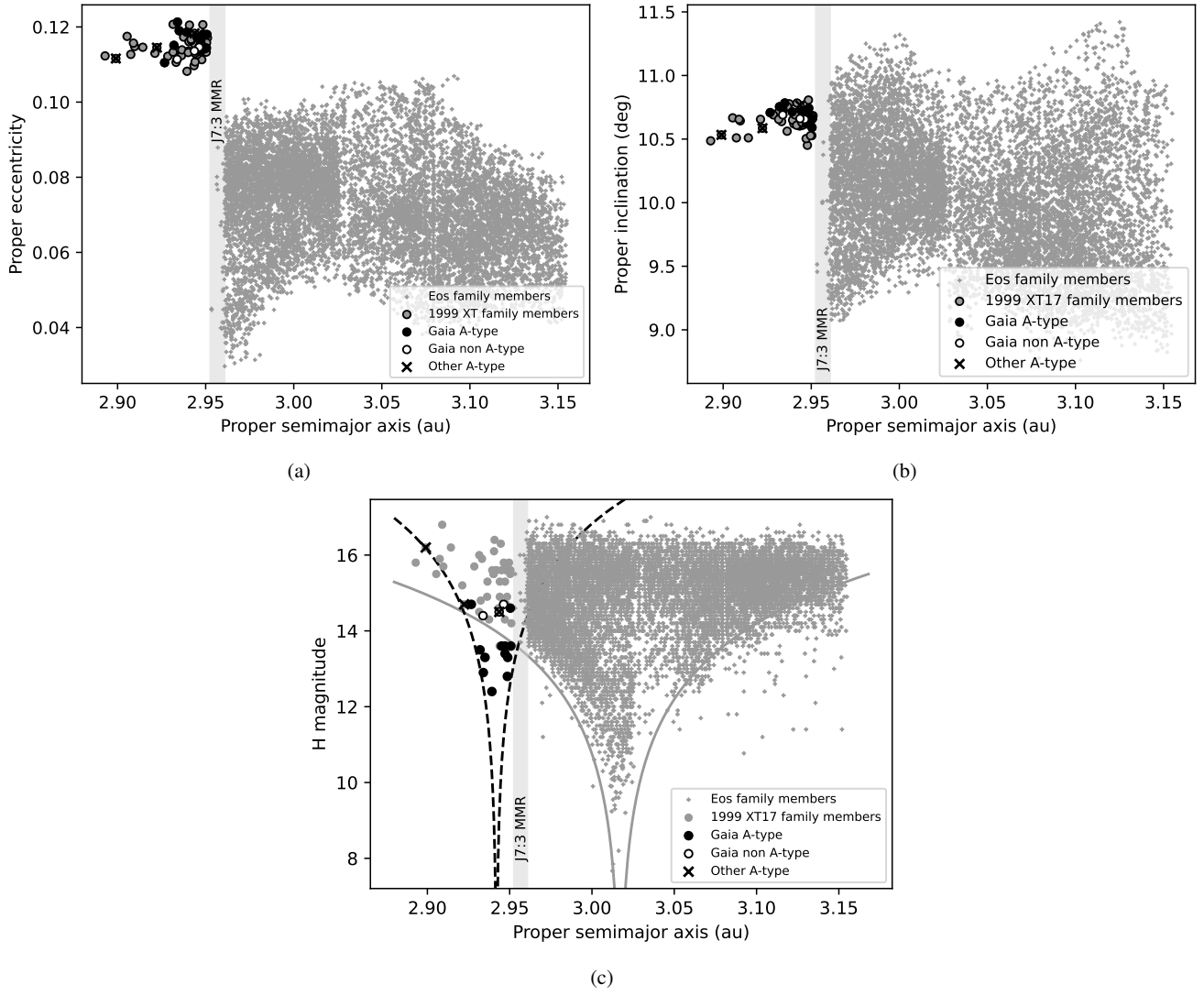


Fig. E.2. Proper orbital elements (panel a and b) and V-shape (panel c) plots of the 1999 XT17 family and the (221) Eos family. The members of Eos are displayed as grey crosses, and those of the 1999 XT17 family as circles. Black filled circles are 1999 XT17 family members that were classified as A-types by our algorithm, while the empty circles are members that were not classified as A-types. Family members that were classified A-types in the literature and have no DR3 spectra are marked as "X". Among the latter group, asteroid 76627 overlaps in the plot with asteroid 88052 (non A-type), due to their similar proper orbital elements and H magnitudes. The position of the 7:3 mean motion resonance with Jupiter is highlighted with a shade of light-grey. The V-shapes of the two families are drawn by eye following the equation $H = 5 \log_{10}(|a - a_c|/C_0)$ with $a_c = 2.942$ au and $C_0 = 1.8 \times 10^{-5}$ au for the 1999 XT17 family, and $a_c = 3.017$ au and $C_0 = 2.25 \times 10^{-4}$ au for the Eos family.