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Environmental effects on satellite galaxies from the perspective of cold gas

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ABSTRACT

Environment plays a pivotal role in shaping the evolution of satellite galaxies. Analysing the properties related to the cold gas phase of satellites provides insights into unravelling the complexity of environmental effects. We use the hydro-dynamical simulations the IllustrisTNG project and the EAGLE simulations, and the semi-analytic models (SAMs) GALaxy Evolution and Assembly (GAEA) and L-Galaxies, in comparison with recent observations from the Westerbork Synthesis Radio Telescope, to investigate the properties of satellite galaxies hosted by haloes with mass $M_{200} > 10^{12.8} M_{\odot}$, and within projected regions ≤ 1.1 virial radius R_{200} . Generally, satellite galaxies selected from semi-analytic models have more HI than those selected from hydro-dynamical simulations across all projected radii, e.g. more than 30 per cent of satellites in the two hydro-simulations are HI depleted, while this fraction is almost zero in SAMs. Furthermore, both hydro-dynamical simulations and SAMs reproduce the observed decrease of HI content and specific star formation rate towards the halo centre. However, the trend is steeper in two hydro-dynamical simulations TNG and EAGLE, resulting in a better agreement with the observational data, especially in more massive haloes. By comparing the two version of GAEA, we find that the inclusion of ram-pressure stripping of cold gas significantly improves the predictions on HI fractions. The refined hot gas stripping method employed in one of the two L-Galaxies models also yields improved results.

Key words: methods: numerical – galaxies: evolution.

1 INTRODUCTION

Interactions with environment influence the evolution of a galaxy. In dense environments, galaxies transform from star forming to passive more efficiently, causing them to be redder than their counterparts in low-density regions (Dressler 1984; Peng et al. 2010). Environmental effects also play a significant role in the evolution of satellite galaxies (van den Bosch et al. 2008; Knobel et al. 2013), that are hosted by haloes more massive than field central galaxies of similar stellar mass. Notably, the quenched fraction of satellite galaxies is significantly higher than that of central galaxies at fixed stellar mass (Wetzel, Tinker & Conroy 2012). Furthermore, as these satellite galaxies approach the centre of their host haloes, their quenched fraction increases (Bluck et al. 2016; Wang et al. 2018a; Hu et al. 2021).

Other properties are also found to depend on the environments. Observational results show that galaxies residing in the denser regions, more massive haloes, or in the inner regions of haloes, tend to have less HI gas (Gavazzi et al. 2006; Cortese et al. 2011;

Fabello et al. 2012; Brown et al. 2017). Additionally, the scaling relation between HI and stellar mass is found to be correlated with the colours (g–r) or (NUV–r) (Cortese et al. 2011; Zu 2020), and reveals that red galaxies typically exhibit a lower HI fraction compared to their blue counterparts. Moreover, among star-forming galaxies, an anticorrelation has been found between HI mass and gas-phase metallicity at fixed stellar mass (Bothwell et al. 2013; Hughes et al. 2013; Brown et al. 2018). The evolution of physical properties of galaxies can be used as diagnostics of quenching mechanisms. For instance, Peng, Maiolino & Cochrane (2015) suggested that strangulation would cause an increase in stellar metallicity for satellite galaxies. The rapid gas removal can be ascribed to ram-pressure stripping (RPS) of cold gas (Fabello et al. 2012; Vollmer et al. 2012; Odekon et al. 2016). Observed phenomena including elongated multiphase gas tails and perturbed morphology of gas discs are inevitably attributed to the effect of ram-pressure stripping (Boselli, Fossati & Sun 2022, and references therein).

Generally, environmental effects include: (a) starvation and strangulation, that suppresses the replenishment of the interstellar medium (ISM) by halting accretion of primordial gas and stripping off the warm/hot circumgalactic medium, (Larson, Tinsley & Caldwell

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1980; Balogh & Morris 2000); (b) galaxy mergers, that can ignite star-bursts and consume most of the ISM (Cox et al. 2006; Hopkins et al. 2013); (c) stripping of ISM, i.e. direct removal of the ISM by RPS (Gunn & Gott 1972) and tidal stripping (Dressler & Gunn 1983). These processes are stronger in more massive haloes (Xie et al. 2020). In addition, internal mechanisms, like stellar winds and supernova explosions or active galactic nuclei (AGN) feedback (Bahé & McCarthy 2015; Zinger et al. 2020), can heat or eject gas and make it more likely to be stripped off. Therefore, these feedback processes would be more effective in dense environments.

Quantifying the relative contributions of different processes is challenging. All the above mechanisms are expected to be involved in the quenching process. A satellite galaxy could experience a complex environmental history, with different physical mechanisms playing the dominant role at different stages of its lifetime (De Lucia et al. 2012; Marasco et al. 2016). The impact of these mechanisms should be considered throughout the entire evolution history of galaxies. From the observational point of view, it is impractical to trace a galaxy's evolution and pinpoint the time intervals when different environmental processes are at play. Yet, the radial distribution of galaxies offers a glimpse of the evolution of satellite galaxies falling into massive haloes (Woo et al. 2013). Interpreting such trends is intricate, as the dependence on radial distance is a combined effect of environment density and accretion time (Gao et al. 2004; De Lucia et al. 2012; van den Bosch et al. 2016). Theoretical models of galaxy formation, fully embedded in a cosmological context, are necessary to disentangle these two contributions, as well as to quantify the impact of projection effects.

Numerical tools have been extensively used and developed to improve the modelling of satellite galaxies. A consensus has emerged regarding the dominant quenching mechanisms: the feedback from the central massive black hole to quench massive galaxies (Croton et al. 2006; Somerville & Davé 2015; Beckmann et al. 2017; Donnari et al. 2021a), environmental effects to quench lower-mass galaxies (Henriques et al. 2015; Stevens & Brown 2017; Hough et al. 2023), and RPS to remove cold gas (Brüggen & De Lucia 2008; Stevens et al. 2019). Many works have been dedicated to comparing different hydro-dynamical simulations and semi-analytic models (SAMs; Brown et al. 2017; Shi et al. 2022; Kukstas et al. 2023). The modelling adopted, however, can vary significantly. Semi-analytic models rely on analytic equations to deal with environmental effects, while these are automatically included in hydro-dynamical simulations (Steinhauser, Schindler & Springel 2016; Kulier et al. 2023). While most SAMs only consider environmental effects on satellite galaxies, they overlook the impacts on central galaxies, which can also undergo significant gas stripping, as modelled by Ayromlou et al. (2019). Wang et al. (2018a) argued that this assumption leads to an overestimation of the difference between central and satellite galaxies. SAMs that have implemented both environmental effects on hot and cold gas find good consistency with observed quenched fractions (Cora et al. 2018; Xie et al. 2020). In contrast, hydro-dynamical simulations face challenges in reproducing quenched fractions of low-mass galaxies within cluster haloes (Bahé et al. 2017; Donnari et al. 2019). Hydro-dynamical simulations offer more natural approaches to treat environmental effects, though medium/large cosmological simulations lack the accuracy to account for the details (Tonnesen & Bryan 2021). Moreover, the limited resolution could cause spurious numerical effects, i.e. artificially increase tidal effects and ram-pressure. In hydro-simulations where centrals and satellites are influenced by dense environments equally, the stripping of gas contents starts beyond the virial radius for galaxies falling into cluster haloes

(Zinger et al. 2018). The star formation rates and gas contents of satellite galaxies are suppressed at larger halo-centric distances in the framework of hydro-simulations with respect to semi-analytic models (Ayromlou et al. 2021a).

In this work, we make use of the observational data sample from Hu et al. (2021) to examine satellite galaxy properties as a function of distance to the host halo centre. We then conduct a comparative analysis involving several state-of-the-art hydro-dynamical simulations and semi-analytic models. Our primary goal is to assess the extent to which current theoretical models can accurately reproduce the observed trends. Additionally, we also divide our samples into two groups based on different halo masses, which enables us to explore the influence of halo mass dependence on the radial distribution of satellite galaxies. Our analysis is confined to the region within $1.1 R_{200}$, thereby minimizing effects from nearby neighbour haloes and concentrating solely on the influence of the host halo environment.

This paper is organized as follows: Section 2 presents the observational data we use. In Section 3, we give a detailed description of simulation data, galaxy properties needed in the following discussion, and how we select our galaxy samples in order to have a fair comparison with observation. Key results are discussed in Section 4, and Section 5 presents our discussion and conclusions.

2 OBSERVATIONAL DATA

As described in Hu et al. (2019) and Hu et al. (2021), original H I data come from the Westerbork Synthesis Radio Telescope (WSRT). In total, 351 h of telescope time were used to integrate thirty-six individual pointings in the Sloan Digital Sky Survey (SDSS) South Galactic Cap region ($21 \text{ h} < \text{RA} < 2 \text{ h}$ and $10^\circ < \text{Dec} < 16^\circ$). Data from one pointing were discarded due to bad quality. The half-power beam width of the WSRT is 35 arcmin, and the average synthesized beam size is $108 \text{ arcsec} \times 22 \text{ arcsec}$. The radio astronomy data reduction package MIRIAD (Sault, Teuben & Wright 1995) was used to reduce and calibrate the data. The reduced data cubes have a size of 601×601 pixels with a pixel size of $3 \text{ arcsec} \times 3 \text{ arcsec}$, and a frequency consisting of 8×20 MHz bands with a resolution of 0.15625 MHz and an overlap of 3 MHz for each band. The final data have an overall frequency ranging from 1.406 to 1.268 GHz, which corresponds to a redshift range of $0.01 < z < 0.12$. Considering stronger radio frequency interference at high redshift, the upper limit is hereafter set to $z = 0.11$.

The SDSS DR7 optical catalogue (Abazajian et al. 2009) is then utilized to get accurate redshifts and spatial positions of galaxies covered by radio data for the subsequent stacking procedure. By cross-matching the radio data to the SDSS catalogue, a sample of 1895 galaxies in the redshift bin $0.01 < z < 0.11$ is obtained. This sample is also constrained by the SDSS magnitude limit, with r -band magnitude of galaxies $M_r < 17.7$. The stacking techniques can then be applied to the radio data and a detailed description is given in Hu et al. (2019). The main procedure includes extracting the radio spectra of galaxies, removing the residual continuum emission and de-redshifting the spectra, conserving H I flux density, deriving mass spectra and weights, correcting for beam confusion, and finally obtaining the weighted average of the mass spectra. The integrated H I mass of a stack is calculated by integrating the mass spectrum in the frequency range $[\nu_0 - \Delta\nu, \nu_0 + \Delta\nu]$, where ν_0 is 1420.406 MHz and $\Delta\nu = 1.5$ MHz.

Group classification is obtained from the group catalogue of Yang et al. (2012) based on the SDSS DR7 main galaxy sample in the redshift range $0.01 < z < 0.2$. This catalogue uses iterative methods to identify groups: first, galaxies are assigned to tentative groups

Table 1. Table of ram-pressure stripping and tidal stripping mechanisms treatments in different SAMs.

Ram-pressure stripping				
	G-X20GRA	G-X20RPS	L-H15	L-A21
Gas type	Hot gas	Hot gas and cold gas	Hot gas	Hot gas
Limit	Satellite galaxies	Cold gas stripped only when its radius is larger than hot gas; satellite galaxies	Satellite galaxies within $1R_{200}$; halo mass $M_{200} > 10^{14} M_{\odot}$;	All satellite and central galaxies
Tidal stripping				
Gas type	Hot gas	Hot gas	Hot gas	Hot gas
Limit	All FOF satellites	All FOF satellites	Satellites within $1R_{200}$ of all haloes	All FOF satellites

using the Friends of Friends (FOF; Davis et al. 1985) algorithm; then, the characteristic luminosity of each tentative group is determined; next, the mass, size, and velocity dispersion are estimated for the dark matter halo associated with each tentative group; then galaxies are then re-assigned using halo information; the steps described above are reiterated until the group memberships converge. The final halo masses are obtained using the halo mass function given by Warren et al. (2006). Satellite galaxies located within R_{180} of a halo and have stellar masses between (10^{10} , $10^{11.5} M_{\odot}$) are used. Under these restrictions, our final sample includes 457 galaxies from a total of 115 haloes [halo mass range ($10^{11.6}$, $10^{14.7} M_{\odot}$)] among which, there are 54 star-forming galaxies.

The following analysis will include several properties of galaxies and the derivation of our data is summarized below. Stellar masses, specific star formation rates (sSFR) and gas-phase metallicities are extracted from the MPA-JHU (Max – Planck Institute for Astrophysics – John Hopkins University) value-added galaxy catalogue of SDSS DR7 (Kauffmann et al. 2003). Gas-phase metallicities can only be calculated for star-forming galaxies due to the emission limit of the OH line, so we only use star-forming galaxies to study gas-phase metallicity in the following analysis. As for stellar metallicity, we use the catalogue based on SDSS DR2 (Gallazzi et al. 2005). The cross-matching of the radio data and SDSS DR2 excludes data from 4 pointings due to the incompleteness of SDSS DR2 compared to SDSS DR7. However, this exclusion has little effect on our final results. Under the same restrictions described above, the final sample for SDSS DR2 counterparts has 352 galaxies from 94 haloes. We only use this sample for stellar metallicity analysis.

3 SIMULATION DATA

In this section, we outline the two semi-analytical galaxy formation models and two hydro-simulations used in this work. Table 1 lists the treatments of environmental effects, specifically ram-pressure stripping and tidal stripping in each semi-analytic model.

3.1 GAEA

We select two versions of the model GALaxy Evolution and Assembly (GAEA) semi-analytic model from (Xie et al. 2020, hereafter X20). X20 is developed on the basis of De Lucia & Blaizot (2007), with a treatment for the non-instantaneous recycling of metals, energy, and gas (De Lucia et al. 2014), a modified stellar feedback method (Hirschmann, De Lucia & Fontanot 2016), an explicit modelling of H_2 -based star formation and partitioning of $H I$ and H_2 (Xie et al. 2017, hereafter X17). X20 updated the treatment of environmental

effects and hot gas and cold gas step-by-step and constructed three updated versions. Here, we use two versions of the model called GRADHOT (gradual stripping of hot gas) and RPSCOLD (ram-pressure stripping of the cold gas) in X20. We refer to them in the following as G-X20GRA and G-X20RPS, respectively.

The G-X20GRA model implements a treatment for tidal stripping and RPS to gradually remove hot gas from satellite galaxies, while the G-X20RPS model additionally includes RPS of cold gas. The tidal stripping of hot gas from satellite galaxies is assumed to be proportional to the stripping of dark matter from the parent dark matter substructures. The RPS of hot gas and cold gas is computed by comparing the ram-pressure and gravity-binding pressure. Moreover, the G-X20RPS model assumes cold gas to be protected by a hot gas shell. Only cold gas outside the hot gas radius can be removed by RPS. X20 shows that both the G-X20GRA and G-X20RPS are in good agreement with quenched fractions measured for SDSS galaxies. The G-X20RPS can reproduce the gas scaling relations measured for central and satellite galaxies from xGASS (the extended GALEX Arecibo SDSS Survey; Catinella et al. 2018) and xCOLDGASS (the extended CO Legacy Database for GASS; Saintonge et al. 2017). GAEA is also able to reproduce the observed evolution of $z < 0.7$ stellar-phase metallicity and the evolution of gas-phase metallicity at $z < 3.5$ (Fontanot et al. 2021), taking into account the uncertainties in observations.

Both models are based on the Millennium Simulation (MS; Springel et al. 2005), in which a WMAP1 (first-year Wilkinson Microwave Anisotropy Probe) cosmology is adopted ($\Omega_m = 0.25$, $\Omega_{\Lambda} = 0.75$, $\Omega_b = 0.045$, $h = 0.73$, $\sigma_8 = 0.9$). The Millennium traces 2160^3 particles from $z = 127$ to the present day. The box size is 685 Mpc with particle mass $1.2 \times 10^9 M_{\odot}$.

3.2 L-GALAXIES

We use two versions of the L-GALAXIES model, Henriques et al. (2015, hereafter L-H15) and Ayromlou et al. (2021b, hereafter L-A21), based on the merger trees from the MS but re-scaled to the first-year Planck cosmology (Planck Collaboration XIII 2016) with $\Omega_m = 0.315$, $\Omega_{\Lambda} = 0.685$, $\Omega_b = 0.0487$, $h = 0.673$, $\sigma_8 = 0.829$. The re-scaled MS has a box size of 714 Mpc and a particle mass of $1.43 \times 10^9 M_{\odot}$.

L-H15 considers tidal stripping to gradually remove hot gas for all satellite galaxies within R_{200} . RPS of hot gas is also considered, but only for satellites in cluster haloes $M_{200} > M^{14} M_{\odot}$. L-H15 shows good agreement with the observed red fractions up to $z \sim 3$, and the distribution of specific star formation rate at $z \sim 0$.

The work of [L-A21](#) includes several improvements compared to [L-H15](#), including updates described in [Henriques et al. \(2020\)](#) (hereafter [L-H20](#)). [L-H20](#) is updated for the treatment of gas cooling, the partition of H I and H₂, and H₂-based star formation in annuli ([Fu et al. 2012](#)), as well as non-instantaneous chemical enrichment ([Yates et al. 2013](#)). The re-calibrated [L-H20](#) provides similar red fractions at $z = 0$ and lower red fractions at higher redshifts compared to [L-H15](#).

Based on [L-H20](#), [L-A21](#) includes an updated approach to deal with RPS of hot gas in a more self-consistent manner ([Ayromlou et al. 2019](#)). The algorithm extracts a Local Background Environment (LBE) for each galaxy by analysing the dark matter particles. The ram pressure applied on each galaxy is measured from the of the LBE density and the relative velocity between the galaxy and its LBE. This version applies RPS to all galaxies, including central ones. The model quenched fractions match reasonably well the observations up to $z \sim 2$.

3.3 EAGLE

The EAGLE simulations ([Crain et al. 2015](#); [Schaye et al. 2015](#); [McAlpine et al. 2016](#)) are a suite of hydro-dynamical simulations with cosmological parameters from the Planck Collaboration [XVI \(2014\)](#). ($\Omega_m = 0.307$, $\Omega_\Lambda = 0.693$, $\Omega_b = 0.04825$, $h = 0.6777$, $\sigma_8 = 0.8288$). This simulation is run with a modified version of the GADGET-3 TREESPH code described by [Springel \(2005\)](#). Dark matter haloes are identified by the FOF algorithm and galaxies with the SUBFIND algorithm ([Springel et al. 2001](#); [Dolag et al. 2009](#)). In this work, we use the reference simulation with the largest volume named Ref-L100N1504. Its side length is 100 cMpc. The initial gas-particle mass is $1.81 \times 10^6 M_\odot$ and the dark matter particle mass is $9.70 \times 10^6 M_\odot$.

EAGLE reproduces the observed main sequence at $z = 0.1$ ([Schaye et al. 2015](#)). The quenched fraction of satellite galaxies in massive haloes with $M_h < 10^{14} M_\odot$ is much higher than the quenched fractions measured for SDSS ([Bahé & McCarthy 2015](#)). [Wang et al. \(2018b\)](#), however, found good agreement between EAGLE and SDSS for the quenched fractions against halo mass for central and satellite galaxies separately. The post-processed H I fractions of EAGLE galaxies in different environments are generally consistent with observations ([Marasco et al. 2016](#)). The scaling relations between stellar-phase/gas-phase metallicity and stellar mass of EAGLE agree well with observation for galaxies with $M_* > 10^{10} M_\odot$. There are significant deviations for lower-mass galaxies.

3.4 IllustrisTNG

The IllustrisTNG project ([Marinacci et al. 2018](#); [Naiman et al. 2018](#); [Nelson et al. 2018](#); [Pillepich et al. 2018](#); [Springel et al. 2018](#), hereafter TNG) is a suite of cosmological hydrodynamical galaxy formation simulations run with the moving-mesh code AREPO ([Springel 2010](#)). The initial cosmological parameters used in this simulation are Planck cosmological parameters ($\Omega_m = 0.3089$, $\Omega_\Lambda = 0.6911$, $\Omega_b = 0.0486$, $h = 0.6774$, $\sigma_8 = 0.8159$). In this study, we use the publicly available data from TNG100, which consists of 1820^3 dark matter particles and 1820^3 baryonic particles in a 100 cMpc box (in order to have a fair comparison with EAGLE, we do not use the 300 cMpc box). The dark matter particle mass is $7.5 \times 10^6 M_\odot$. The baryonic particle mass is $1.4 \times 10^6 M_\odot$.

The quenched fractions are in broad consistency with observational measurements ([Donnari et al. 2021b](#)). The post-processed H I reasonably agrees with observational data from ALFALFA and

xGASS ([Stevens et al. 2019](#)). TNG also reproduces observed gas-phase metallicity-stellar mass relations in the redshift range $0 < z < 2$ ([Torrey et al. 2019](#)).

3.5 Galaxy properties

In this work, we investigate environmental effects on satellite galaxies at $z = 0$. More specifically, we compare the H I mass, star formation rate (SFR), and metallicity of satellite galaxies as a function of the projected distance to their host halo centre with observation. For EAGLE, all properties are defined to be the sum of the properties of all star particles that belong to the corresponding subhalo and that are within a 3D aperture with radius 30 pkpc. For TNG, stellar mass, SFR and H I mass are calculated by total properties of all member particle/cells which are bound to this subhalo. Metallicities are calculated by all star or gas particles bound to the parent subhalo but restricted to gas cells within twice the stellar half mass radius.

The stellar mass and SFR are direct outputs of the simulations, all assuming a [Chabrier \(2003\)](#) initial mass function. Galaxies may be characterized by extremely low SFR in hydro-dynamical simulations and SAMs, whose values are not well constrained in observational measurements. These measurements are often assigned a log sSFR around -12 ([Brinchmann et al. 2004](#)) (sSFR is defined as SFR/M_*). For consistency with this choice, we re-assign the sSFR of all model galaxies with $\log \text{sSFR} < -12 \text{ yr}^{-1}$ to random values in the range of $-12.5 < \log \text{sSFR} < -11.5 \text{ yr}^{-1}$. The H I masses and metallicities are computed differently for each model and simulation, as described below.

(i) **H I mass:** G-X20GRA, G-X20RPS, and [L-A21](#) implement an explicit treatment to partition cold gas into H I and H₂. G-X20RPS and G-X20GRA use the empirical relation from [Blitz & Rosolowsky \(2006\)](#) to account for the ratio between H I and H₂. In [L-A21](#), the H₂ fraction is modelled using the fitting equations of [McKee & Krumholz \(2010\)](#) originated from [Krumholz, McKee & Tumlinson \(2009\)](#). EAGLE and TNG provide predictions for post-processed H I. In EAGLE, H I mass is calculated in each gas cell ([Rahmati et al. 2013](#)) using the empirical observational relation ([Blitz & Rosolowsky 2006](#)), in which the fraction of H I depends on the mid-plane pressure of this cell. For TNG, the detailed calculation of H I mass is given by [Diemer et al. \(2018, 2019\)](#). They provide several methods of partitioning H I and H₂. Here, we choose the model named GD14 ([Gnedin & Draine 2014](#)). [L-H15](#) does not provide estimates for H I mass. We use a constant H I gas ratio $M_{\text{HI}}/M_{\text{coldgas}} = 54$ per cent as in [Henriques et al. \(2015\)](#). Here, we assume 75 per cent of cold gas is neutral hydrogen, and 70 per cent of neutral hydrogen is in the atomic phase.

[Fig. 1](#) shows the H I and stellar mass functions of all hydro-dynamical simulations and SAMs at $z = 0$. All hydro-dynamical simulations and SAMs match well the stellar mass function, with the exception of the massive end. Here, however, observational measurements are affected by large uncertainties due to small number statistics/volumes. For the H I mass functions, TNG is in good agreement with observational measurements. SAMs coupled with MS can only resolve galaxies with $M_{\text{HI}} > 10^9 M_\odot$ due to limited resolution. EAGLE predicts a lower H I mass function over the entire mass range, with a bump at around $M_{\text{HI}} \sim 10^8 M_\odot$. According to [Crain et al. \(2017\)](#), this corresponds to galaxies having no stellar particles, and is a result of numerical artificial effects.

(ii) **Metallicity:** The stellar metallicity is presented as Z_*/Z_\odot . For SAMs, Z_* is the total mass of all elements heavier than Helium

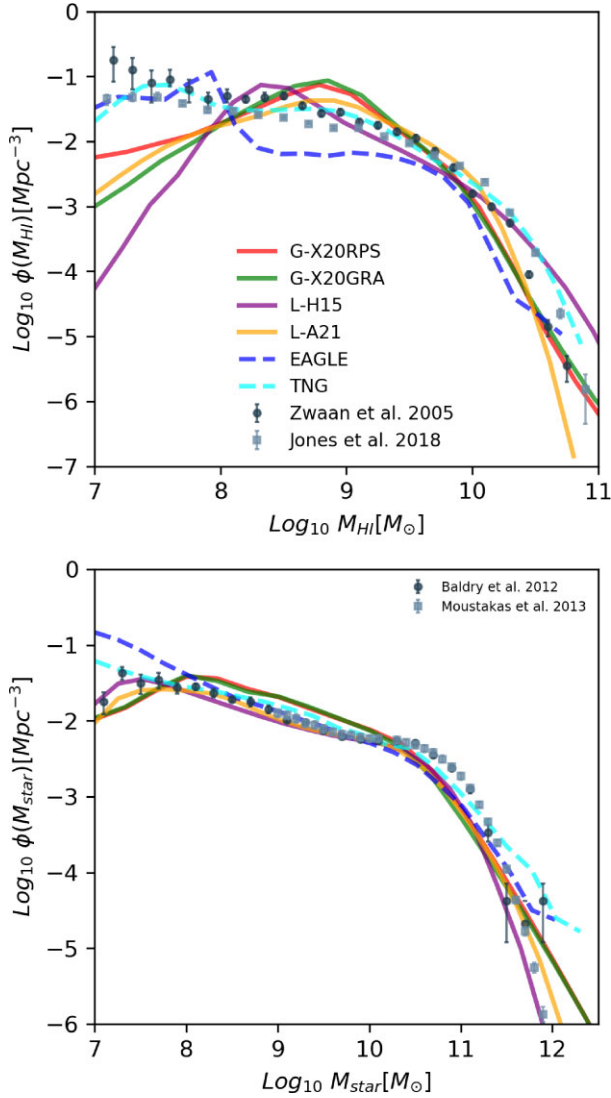


Figure 1. The top and bottom panels show the H I mass functions and the stellar mass functions at $z \sim 0$ for each model/simulation as labelled. H I mass functions are compared to the observations by Zwaan et al. (2005) (circle) and Jones et al. (2018) (square). Stellar mass functions are compared to Baldry et al. (2012) (circle) and Moustakas et al. (2013) (square).

divided by the total stellar mass. For EAGLE and TNG, the stellar metallicity is the total abundance of several elements. The solar metallicity is $Z_{\odot} = 0.02$.

Gas-phase metallicity is represented by oxygen abundance $12 + \text{Log}_{10}(\text{O}/\text{H})$, where O/H is the number ratio of the two particles. G-X20RPS, G-X20GRA, and L-A21 take into account the yields of individual elements, and output masses of oxygen and hydrogen directly. For L-H15, we use the approximation in Ma et al. (2016), which assumes that oxygen abundance scales linearly with the gas mass metallicity (that is, $12 + \text{Log}_{10}(\text{O}/\text{H}) = \text{Log}_{10}(Z_{\text{gas}}/Z_{\odot}) + 9$). Here, Z_{gas} has the same definition as Z_{\star} but in the gas phase. Both EAGLE and TNG provide separate Oxygen and Hydrogen mass fractions for the galaxies in the star-forming gas phase in the download catalogue.

In the observations, the gas-phase metallicity measurements are based on strong emission lines that are only visible for star-forming galaxies. In order to make a fair comparison, we also select star-

forming galaxies with $\text{sSFR} > 10^{-11} \text{ yr}^{-1}$ in the analysis of gas-phase metallicity.

(iii) **Projected distance:** For consistency, we measure projected distances between selected galaxies and the centre of each system in hydro-dynamical simulations and SAMs. In order to account for projection effects, we set x , y , and z directions as line of sight in turns. The projection depth is $10 \text{ Mpc } h^{-1}$. The impact of projection is discussed in Appendix A. The projected distances are normalized by R_{200} , the halo radius where the average enclosed density is 200 times the critical density of the Universe. For the observational data, which use R_{180} as the normalization radius (defined as the halo radius with an average enclosed density 180 times the critical density), we convert it to R_{200} using $R_{180} \approx 1.05R_{200}$ (Reiprich et al. 2013).

3.6 Galaxy sample selection

From each hydro-dynamical simulations and SAMs, we select haloes with $10^{12.8} < M_{200}/M_{\odot} < 10^{14.7}$. For each halo, we select all galaxies with a projected distance to the central galaxy less than $1.1R_{200}$. Following the observational sample, the selection criteria for galaxies from hydro-dynamical simulations and SAMs are $10^{10} < M_{\text{star}}/M_{\odot} < 10^{11.5}$ at $z = 0$. Table 2 lists sample sizes of all SAMs, hydro-dynamical simulations, and observations.

Fig. 2 illustrates the comparison between the halo mass distribution of the observational sample and the distribution obtained from hydro-dynamical simulations and SAMs. The observed galaxies have a minimum host halo mass of $10^{11.6} M_{\odot}$. However, in order to maintain consistency within the halo mass range $M_{200} > 10^{13} M_{\odot}$, we have chosen a higher halo mass cut of $M_{200} > 10^{12.8} M_{\odot}$ for the model galaxies. The selected samples from hydro-dynamical simulations and SAMs display a slightly lower fraction of massive cluster haloes than the observational sample. We will discuss the halo mass dependence in Section 4.2.

Fig. 3 shows the cumulative distribution function (CDF) of galaxy properties for selected samples. The scatters between different hydro-dynamical simulations and SAMs are notable. The observational results fall within the scatter range. In general, galaxies selected from hydro-dynamical simulations and SAMs are less massive, more star forming, and less enriched than those in the observational sample. It is worth noting that the observational sample we used is biased to lower SFR, compared to the SDSS. The two GAEA samples show higher sSFR distribution than other samples because the quenched cloud in GAEA has a peak sSFR located at $\sim 10^{-11} \text{ yr}^{-1}$ (Xie et al. 2020).

The scatter between hydro-dynamical simulations and SAMs is even larger for the H I mass distribution. Generally, galaxies selected from semi-analytic models have more H I than galaxy samples from hydro-dynamical simulations. In EAGLE and TNG, there are ~ 38 per cent and ~ 45 per cent of galaxies contain no H I at all. In semi-analytic models, however, even quenched galaxies can have a H I mass that is larger than $\sim 10^8 M_{\odot}$ in L-H15 models and G-X20GRA models. Galaxies in the G-X20RPS and the L-A21 models have H I masses down to $\sim 10^7 M_{\odot}$. The H I mass distribution of the observational sample is not shown in this panel, as we use the stacked H I mass within a specific region, thus lack the H I mass for individual galaxies. We understand that the observed galaxies with extremely low H I mass are also included in the analysis. Therefore, we also include simulated galaxies with no H I content.

Table 2. Number of galaxies and haloes in each selected samples. Rows marked with projected/3D are selected using projected/3D distances within $1.1 R_{200}$. Galaxies are projected along the x , y , and z directions in turns. Therefore, the sample size is tripled when selected by projected distances.

	G-X20RPS	G-X20GRA	L-H15	L-A21	EAGLE	TNG	Observation
Galaxy number	1034 955	965 878	652 352	505 027	2557	4712	457
Star-forming galaxy number	691 788	735 318	288 059	189 272	993	1530	54
Star-forming galaxy fraction	0.668	0.761	0.441	0.375	0.388	0.325	0.118
Halo number	199 905	193 632	188 242	175 973	565	792	115

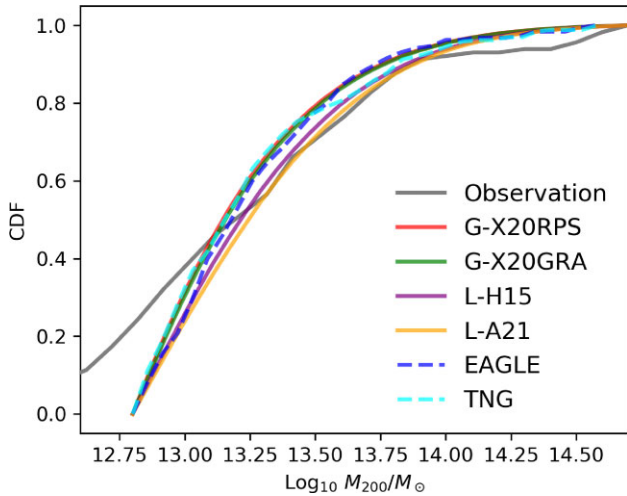


Figure 2. Cumulative distribution functions (CDFs) of host halo mass. The y -axis value corresponds to the fraction of haloes with mass smaller than M_{200} .

4 RESULTS

In this section, we compare the entire model selected samples with observation, then divide them by host halo mass into two groups to investigate the halo mass dependence. Additionally, we extend the previous observational survey to explore the metallicity trends. Given the large scatter across various samples, we focus on the trend as a function of projected distances, rather than on the normalization.

4.1 H I mass and star formation rate

The observational work of Hu et al. (2021) has confirmed a decreasing trend of H I mass for satellite galaxies from the outside to the halo centre. Here, we discuss whether the trend exists in hydro-dynamical simulations and SAMs, and how this depends on the star formation activity of galaxies, as shown in Fig. 4. For consistency with observation, we plot the logarithmic mean value calculated by using jackknife resampling method for SAMs/hydro-dynamical simulations. Additionally, we also plot the median values, which are less biased and better reflect the intrinsic trends than the logarithmic mean values. We plot the median properties for only three SAMs, G-X20RPS, EAGLE, and TNG, to avoid overcrowding the figure with too many lines.

The H I mass and star formation rate correlate primarily to stellar mass. Therefore, it is important to check if there’s any systematic bias between the stellar mass of observations and simulations at given projected halo-centric distances. The upper left panel of Fig. 4 displays the mean stellar masses of galaxies at fixed radial

distances. The differences between the observation and simulations are small. Also, the mean stellar mass of the observational sample does not depend on the radial distance of galaxies significantly. The same trend is found for all hydro-dynamical simulations and SAMs.

The bottom two panels of Fig. 4 show the H I masses and H I mass fractions as a function of their distance to the halo centre. All hydro-dynamical simulations and SAMs predict decreasing H I mass and H I mass fraction with decreasing distances, which is consistent with observations. The decline of H I mass and fractions becomes steeper in the inner regions, indicating a stronger impact of the environment in the halo centre.

Comparing the G-X20RPS and G-X20GRA, the former sample shows a steeper decline of H I than the latter. The RPS of cold gas is essential to reproduce the evolution trend observed in data in the GAEA model. Fig. 5 shows CDFs of H I mass in narrow bins of projected distances. With the additional implementation of RPS on cold gas, the G-X20RPS sample includes a notably larger fraction of low H I mass galaxies than the G-X20GRA. The difference becomes more obvious at smaller distances.

In the case of L-Galaxies, the L-A21 sample is in better agreement with the observation than L-H15 in Fig. 4. Both models do not consider the RPS of cold gas. The decrease in H I is a reflection of the decrease of hot gas and hence of the suppression of gas cooling. Ayromlou et al. (2021b) found that the hot gas to stellar mass ratio is lower in the L-A21 for $\sim 10^{10.5} M_{\odot}$ galaxies in group haloes ($\sim 10^{13.5} M_{\odot}$), with respect to L-H15. As a result, galaxies from L-A21 retain less cold gas than those from L-H15. As shown in Fig. 5, the L-A21 sample includes a non-negligible fraction of H I-poor galaxies at $M_{\text{HI}} < 10^8 M_{\odot}$. This fraction increases towards the halo centre, where such H I-poor galaxies are not found in the L-H15.

Both the RPS on cold gas in G-X20RPS and the updated treatment of RPS stripping on hot gas in L-A21 help to improve the model predictions with respect to previous renditions of the same physical models.

As for hydro-dynamical simulations, both EAGLE and TNG predict a declining trend for H I mass and fractions that is similar to the observations (see Fig. 4). In Fig. 5, the cumulative H I mass distributions of galaxies in TNG and EAGLE are biased towards low H I masses, compared to the semi-analytic models. Even at the virial radius, more than 10 per cent of galaxies in hydro-simulations have no H I. Moving closer to the halo centre, more galaxies in hydro-simulations have zero H I mass. This indicates a strong suppression of H I in EAGLE and TNG due to environmental effects.

The logarithmic mean values in Fig. 4 are dominated by the H I-rich galaxies, because of the non-normal distribution of H I mass. We also include the median values to provide a less biased trend. Interestingly,

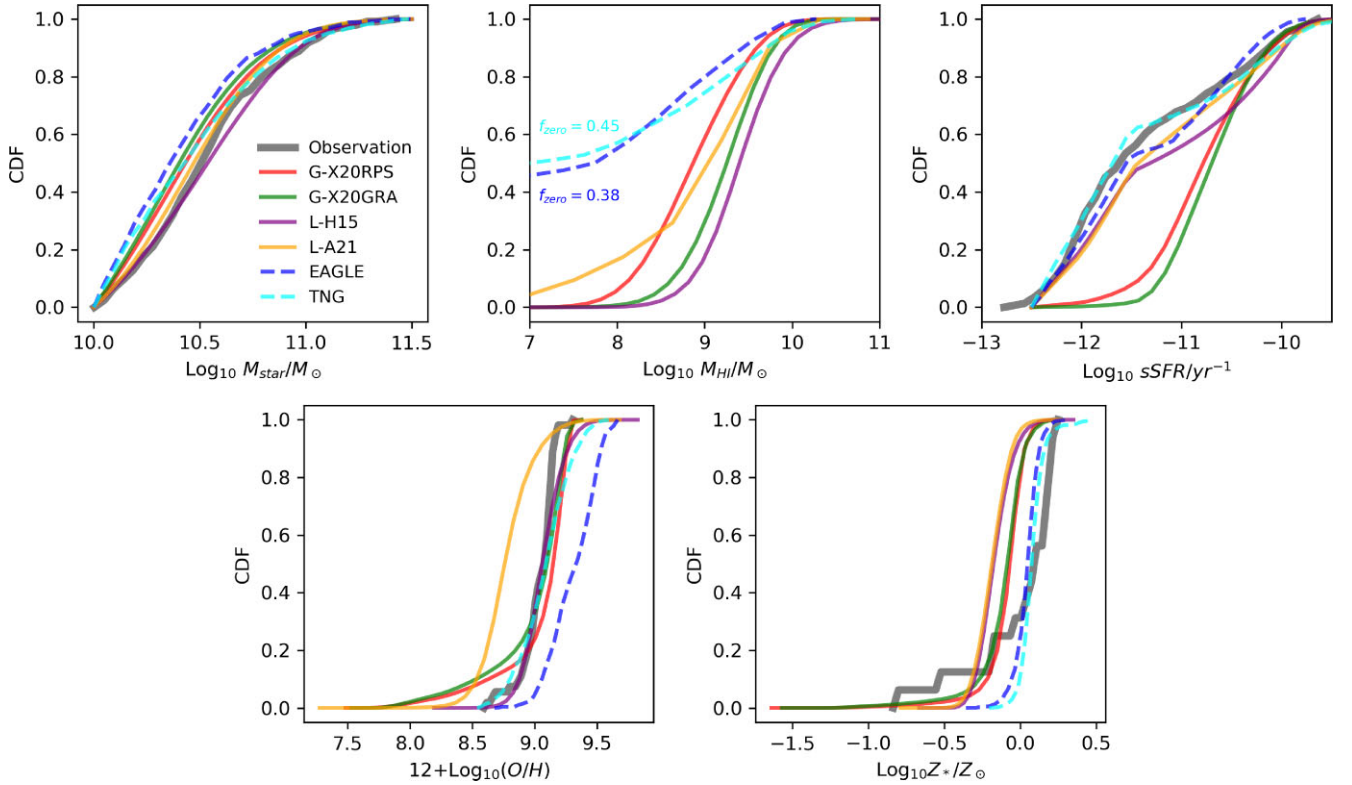


Figure 3. CDFs of stellar mass (upper left), H I mass (upper middle) and sSFR (upper right) as well as of gas-phase metallicity (lower left) and stellar metallicity (lower right). Note that the panel of H I mass does not include the observational sample as the spectral stacking technique used to estimate the observed H I mass does not give H I mass information for individual galaxies. In the panel of H I mass distribution, the coloured numbers represent the fractions of galaxies with zero H I mass in the TNG and EAGLE, respectively. In the panel of sSFR, galaxies in SAMs/hydro-dynamical simulations with $\text{sSFR} < 10^{-12} \text{ yr}^{-1}$ are assigned random values between $(10^{-12.5}, 10^{-11.5}) \text{ yr}^{-1}$. The fractions of galaxies with no star formation are: **L-H15**: 11.42 per cent; **L-A21**: 22.46 per cent; **EAGLE**: 44.94 per cent; **TNG**: 54.50 per cent. The two GAEA versions do not include any zero SFR galaxies. The lower left plot of gas metallicity only displays the distribution of star-forming galaxies ($\text{sSFR} > 10^{-11} \text{ yr}^{-1}$).

while the average H I masses of TNG and EAGLE do not significantly differ from the semi-analytic model G-X20RPS, their median values are offset significantly below the average trends. In the EAGLE and TNG simulations, the median H I fractions remain below $10^{-1.5}$ even at the virial radius. The fact that the mean and median are so different is a consequence of having a lot of galaxies with zero H I mass, as noticed in Fig. 5. The suppression of H I for galaxies in dense environments is much stronger in hydro-simulations than in semi-analytic models. Surveys for H I-poor galaxies around cluster haloes may provide more constraints on the environmental quenching physics.

The sSFR as a function of the projected radius is shown in the upper-right panel of Fig. 4. All hydro-dynamical simulations and SAMs except EAGLE samples have higher or comparable average sSFRs than observational samples at all distances. As mentioned before, the observational sample is biased to lower sSFR galaxies. For the observational sample, the average sSFR decreases from $10^{-10.6} \text{ yr}^{-1}$ at $0.7R_{200}$ to 10^{-11} yr^{-1} at $0.2R_{200}$. All hydro-dynamical simulations and SAMs show a similar decreasing trend with the average sSFR declining ~ 0.4 dex from $1R_{200}$ to $0.25R_{200}$.

By comparing the G-X20RPS sample with the G-X20GRA sample, we find that the inclusion of RPS on cold gas slightly amplifies quenching effects within the $0.5R_{200}$. It can also be noticed that the updated RPS approach implemented on **L-A21**, which considers the effect of ram-pressure stripping on central galaxies, intrigues a suppression of star formation before galaxies fall into the host haloes

compared to **L-H15**. The difference between these two models can be witnessed at R_{200} .

The radial distribution of average sSFR for the EAGLE sample matches very well the observations. In addition, its median sSFR drops significantly at around $0.5R_{200}$, as a consequence of the rapid transition of EAGLE galaxies into a quenched phase at this radius. In the case of the TNG sample, the average sSFR of galaxies decreases from the $1R_{200}$ to $0.3R_{200}$, then increases in the innermost bin. The median sSFR exhibits a continuous decrease from R_{200} to the innermost bin. We have checked that the innermost average sSFR is biased by a few galaxies with very high star formation rates in the halo centre. Interestingly, the majority of galaxies in the TNG sample at the virial radius are quenched, when half of the galaxies still have a large fraction of H I $M_{\text{HI}} > 10^9 M_{\odot}$ (Fig. 5).

4.2 Halo mass dependence

It is widely accepted that environmental effects depend on halo mass. In this section, we split galaxies into two groups based on their host halo mass. Specifically, we select galaxies in low-mass haloes with $M_{200}/M_{\odot} < 10^{13.2}$ and cluster haloes with $M_{200}/M_{\odot} > 10^{14.0}$. The results are shown in Fig. 6. Note that hydro-simulations, especially EAGLE, include fewer massive cluster haloes than SAMs and observational data, making their predictions more affected by Poisson noise.

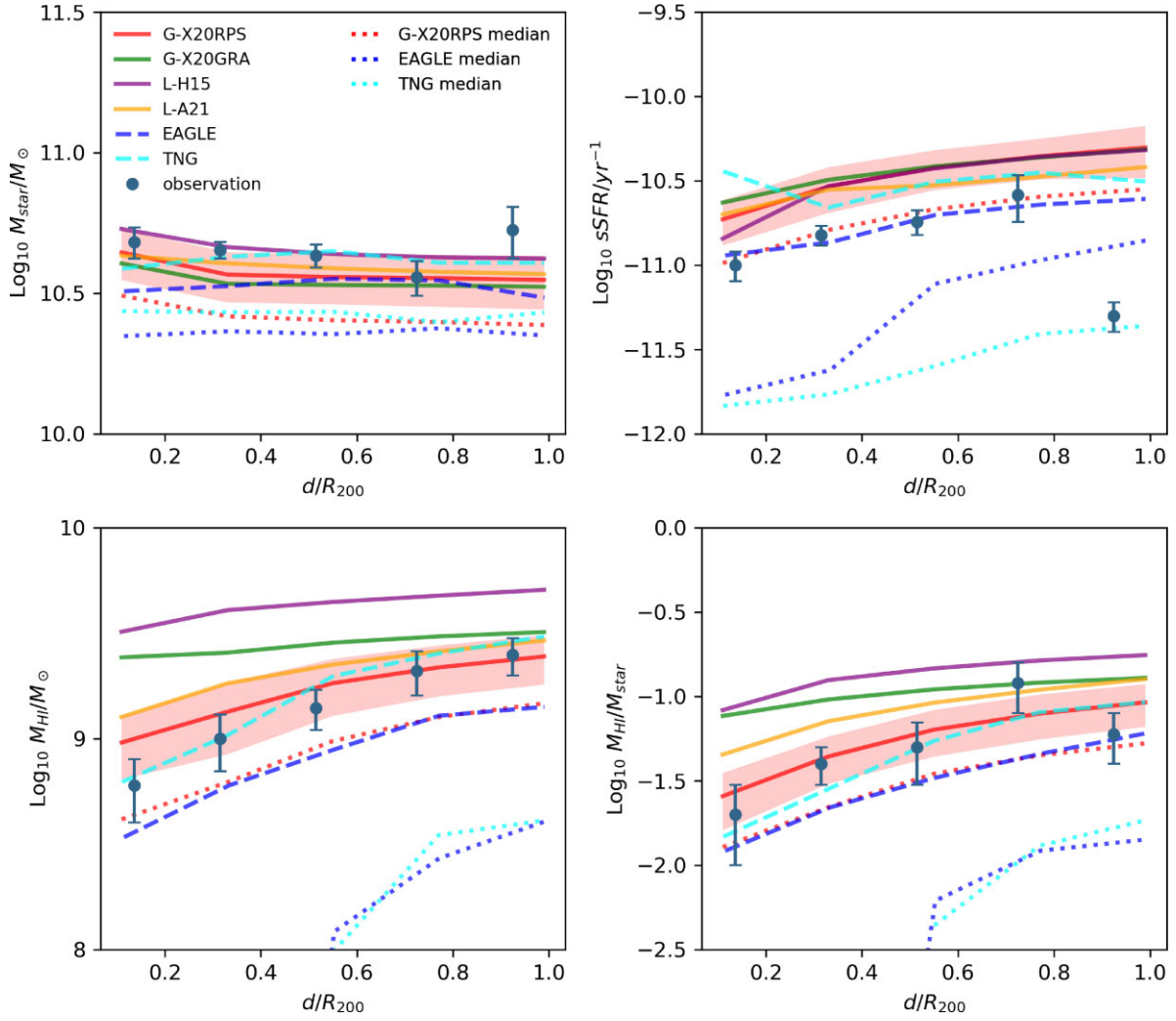


Figure 4. Radial distribution of averaged stellar mass (upper left), sSFR (upper right), H I mass (lower left) and H I mass function (lower right) and their comparison to the observation (symbols). Solid lines and dashed lines represent the average values using the jackknife resampling method, the same as the observation. The shaded area denotes errors and displays only for model G-X20RPS for clarity. To avoid excessive lines, we opt to show only the median distribution as dotted lines for G-X20RPS, EAGLE, and TNG. Note that the bin at the largest distance of the observational sample should be interpreted with caution due to its small sample size (Hu et al. 2021). For example, the lowest value in the last bin of sSFR is caused by galaxies with highest averaged stellar mass.

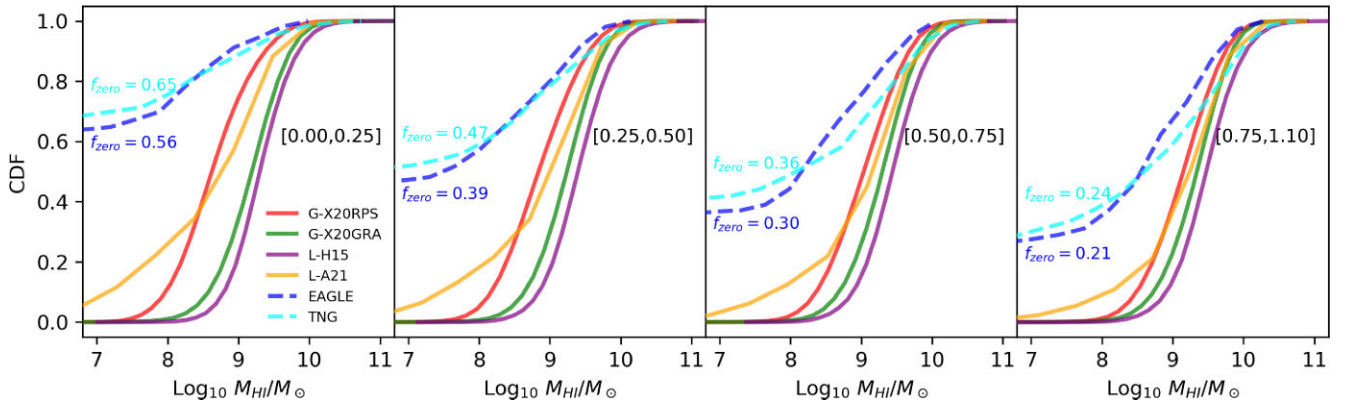


Figure 5. CDFs of H I mass in different normalized radial projected distance bins. Four panels represent ranges (0.00, 0.25), (0.25, 0.50), (0.50, 0.75), (0.75, 1.00), respectively. As in Fig. 3, the coloured percentage represents the fraction of galaxies with zero H I mass of the hydro-dynamical simulations TNG and EAGLE.

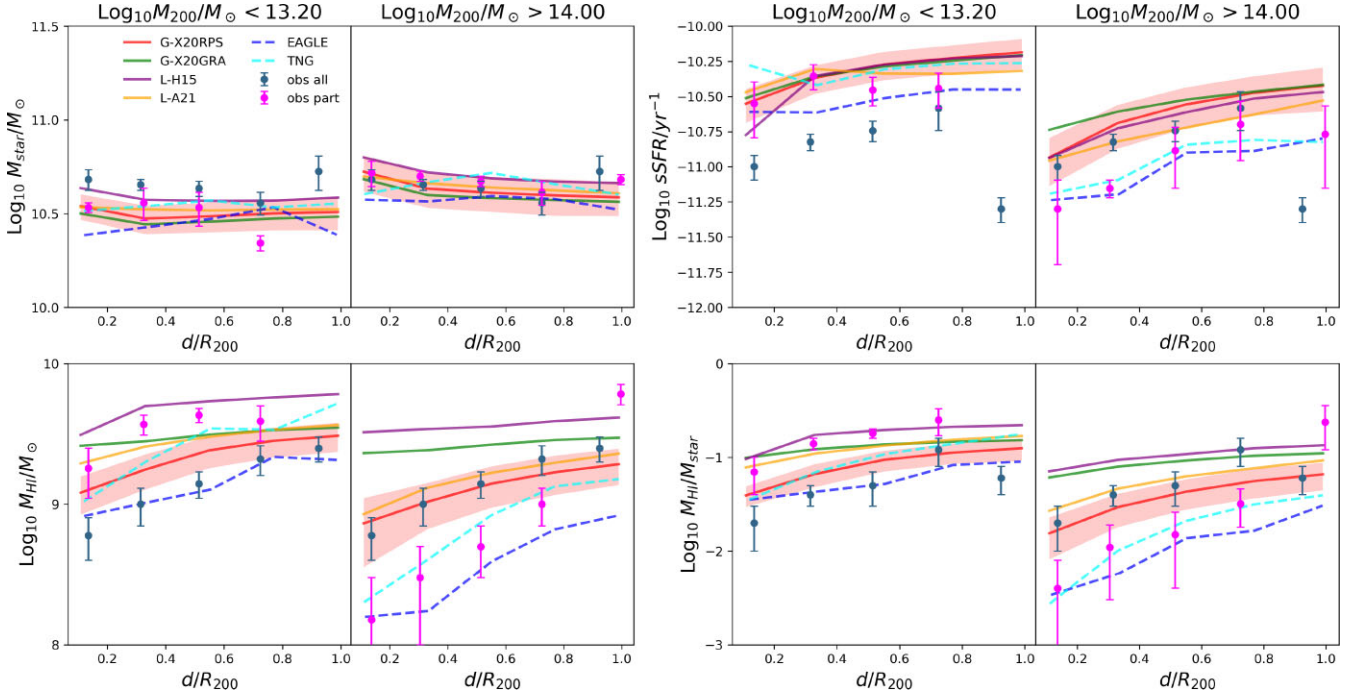


Figure 6. Same as Fig. 4 but in different host halo mass range. From left to right, top to bottom, the paired panels show the radial projected distance distribution of stellar mass, sSFR, H I mass, and H I mass fraction. For each paired panel, the left one shows the results of lower-mass haloes with $M_{200}/M_{\odot} < 10^{13.2}$. The right shows results of more massive haloes with $M_{200}/M_{\odot} > 10^{14.0}$ for simulations as well as the observation. We include the results for observed galaxies in the given halo mass bin and all observed galaxies for reference. Note that the fifth bin (if plotted) of the observation is unreliable due to its small sample size.

The average H I mass and fractions present a clear dependence on the host halo mass across all the samples. Generally, galaxies residing in massive haloes tend to have lower H I mass and fractions compared to those in low-mass haloes. These differences already exist in the outermost bin. At $\sim 0.7R_{200}$ in the observation, the H I fraction in massive haloes is ~ 0.8 dex lower than in low-mass haloes. This decrease is ~ 0.2 dex for all SAMs and ~ 0.4 dex for the two hydro-simulations. Therefore, EAGLE and TNG can better reproduce this decline.

The declining rate of H I also varies depending on the host halo mass. Observed galaxies in the innermost bin have ~ 0.6 dex lower H I mass than galaxies at distances $0.6 < d/R_{200} < 0.8$ in low-mass haloes. This difference becomes ~ 1 dex in massive haloes. This halo mass dependence pattern is only evident in G-X20RPS, L-A21, EAGLE, and TNG. Hydro-simulations better capture the influence of halo mass on H I content. TNG yields the closest agreement with observational results in both low-mass and high-mass haloes. EAGLE reproduces well the data for the most massive haloes, but the predicted H I mass decline towards the centre is milder and its overall H I content is considerably lower than the observations in low-mass haloes. In the framework of G-X20RPS and L-A21, satellite galaxies in massive haloes show higher H I fractions and a milder decrease from the halo edge to the halo centre compared to the observational sample, indicating that the depletion of H I in SAMs are not as efficient as observations suggest. For satellites in low-mass haloes, G-X20RPS predicts a lower H I fraction but a similar observed decreasing pattern from the halo outskirts to the centre, while L-A21 is closer to the observation with a milder decrease. In both cases, L-A21 is ~ 0.1 dex higher than G-X20RPS.

The average sSFR shows similar halo mass dependences. The overall observed sSFR is higher with a milder declining rate in low-mass haloes than in more massive haloes. The same tendency

also holds true for most hydro-dynamical simulations and SAMs. This means that satellite galaxies located within cluster haloes tend to exhibit reduced star-formation activity and begin quenching at greater radii compared to their counterparts in lower-mass haloes. The model L-A21 provides the closest agreement with the observational results for low-mass haloes, while hydro-simulations show better consistency with the observation in high-mass haloes.

It is well known that the environmental effects would deplete the diffuse H I earlier than the more dense star-forming gas. The quenching of satellite galaxies occurs later than the depletion of their H I gas. This ‘delayed quenching’ is found for observed samples in both low-mass and high-mass haloes.

Focusing on G-X20RPS and G-X20GRA, the former model predicts lower H I fractions and sSFR along with steeper decreasing trends towards the halo centre in both low- and high-mass haloes. However, in high-mass haloes, the offset between these two models becomes larger, particularly towards the halo centre, indicating a stronger halo mass dependence in G-X20RPS. This means that RPS of cold gas in the GAEA model is more effective in depleting H I and quenching galaxies in high-mass haloes, which is consistent with the observations.

L-A21 also shows lower H I fractions in low- and high-mass haloes than L-H15. The discrepancy between these two models also widens in higher-mass haloes, suggesting that the updated stripping algorithm implemented in L-A21 more effectively depletes H I in high-mass haloes. However, their sSFR comparison reveals a different behaviour. Though L-A21 generally predicts lower sSFR and is more consistent with the observation, L-H15 predicts a steeper declining trend in both low- and high-mass haloes, providing a more accurate treatment of quenching rates in satellite galaxies.

Though SAMs do show some degree of halo mass dependence, this dependence is not as dramatic as in hydro-simulations. This might be because hydro-simulations have a better treatment of the spatial distribution of the gas components, resulting in stronger HI depletion in high-mass halo mass and stronger quenching mechanisms.

4.3 Environmental effect on metallicity

The gas-phase and stellar-phase metallicity depend on the local density surrounding the galaxy (Peng, Maiolino & Cochrane 2015). Except for the impact of assembly bias, environmental effects on galaxies are also reflected in metallicities. Star formation, gas accretion, and gas stripping could affect the gas-phase metallicity by reducing/increasing the gas fraction or injecting metals into the interstellar medium. The stellar metallicity reflects the evolution history of a galaxy on a longer time-scale. In this section, we compare the radial distribution of gas-phase and stellar-phase metallicities between the observed sample and simulations, as shown in Fig. 7. Since in the MPA-JHU value-added galaxy catalogue, gas-phase metallicity is typically estimated using strong emission lines. For consistency with observations, in this section we consider only star-forming galaxies with $sSFR > 10^{-11} \text{ yr}^{-1}$ from hydro-dynamical simulations and SAMs.

We noticed that the gas-phase metallicity of L-A21 is ~ 0.2 dex lower than the observations. This deviation agrees with Fig. 6 of Yates et al. (2021), where for stellar masses above $10^{10} M_{\odot}$, the gas-phase metallicity typically falls within the range of 8.6 to 8.8, ~ 0.2 dex lower than the previous L-Galaxies model. Conversely, EAGLE predicts metallicities ~ 0.2 dex larger than the observations (see also Bahé et al. 2017).

Regarding stellar metallicity, there is an underestimation of 0.1 or 0.2 dex for all SAMs compared to the observational data, though they are still within the 1 sigma scatter of the observation (Hirschmann, De Lucia & Fontanot 2016; Yates et al. 2021). This discrepancy also arises from our sample selection with $M_{\text{start}} > 10^{10} M_{\odot}$. Therefore, in the following, we focus on the radial trends, and pay less attention to the absolute metallicity values.

The average metallicity of the observational sample does not depend on halo-centric radii at large distances, with a small decrease observed in the innermost radial bin. Interestingly, the stellar-phase metallicity increases towards the halo centre, whereas the gas-phase metallicity decreases. We interpret the trend as a result of multiple reasons. By selecting star-forming galaxies, we tend to select galaxies that either suffer weaker environmental effects and remain star forming or those outside of the host haloes and are selected by projection. The average gas-phase metallicity could also be biased by a few low-metal star-forming galaxies. These influences are witnessed in the results of SAMs.

The evolution trend of the L-Galaxies series model are consistent with observation, despite the fact that galaxies in L-A21 are metal-poorer than the observational sample. Both L-H15 and L-A21 show decreasing gas-phase metallicity and increasing stellar metallicity towards halo centre. However, by tracing back the evolution of all satellite galaxies, we find increasing trend for both stellar and gas-phase metallicity. About half satellite galaxies in the halo centre, however, are excluded in the analysis of gas-phase metallicity because they are quenched. Therefore, the average gas-phase metallicity decreases.

G-X20RPS and G-X20GRA show similar trends in the evolution of gas-phase metallicity, which is consistent with observational results. Their average stellar metallicity of the entire sample also

decreases towards the halo centre. The decreasing trend is caused by satellite galaxies that have experienced recent gas cooling. A few satellite galaxies have a large fraction of hot gas at the time of infall which maintains gas cooling and star formation even when they are satellites. With the replenishment of metal-poor cooling gas, these satellite galaxies are relatively metal-poorer in both the gas-phase and stellar-phase than others with relatively lower star formation rates. These gas-poor galaxies cause a slight downward in the mean metallicities. But the impact on median metallicities is negligible. The median gas-phase metallicity and stellar-phase metallicity are flat or increasing towards the halo centre. We notice that with the RPS of cold gas implemented, the average gas-phase metallicities of galaxies in G-X20RPS are higher than their analogues in G-X20GRA.

Both EAGLE and TNG predict increasing gas-phase metallicity and stellar-phase metallicity towards the halo centre. The median results are consistent with the average values. environmental dependencies of metallicity have been discussed in previous works. Wang, Wang & Chen (2023) found that satellite galaxies are more metal-enriched in more massive cluster haloes, whereas central galaxies are more metal-poor in more massive haloes. They explain the difference as a result of the influence of environmental effects on satellite galaxies, and AGN feedback on central galaxies. Donnan, Tojeiro & Kraljic (2022) have discussed the dependence of gas-phase metallicities on environment for the TNG simulation. They interpret the dependences found as a combined result of halo assembly bias and enriched gas inflow in dense environments. We argue that environmental stripping could increase the metallicity more efficiently in hydro-dynamical simulations than in the semi-analytic models. As stripping occurs outside-in, it would remove metal-poor gas. In semi-analytic models, gas stripping does not affect its metallicity.

5 DISCUSSION AND CONCLUSION

In this paper, we investigate the influence of environment by focusing on the cold gas content of satellite galaxies. In particular, we compare the observed radial distributions of HI mass/fraction, sSFR, gas-phase metallicity, and stellar-phase metallicity with several theoretical models: including the hydro-dynamical simulations Illustris TNG and Eagle, and the semi-analytic models GAEA and L-Galaxies. Additionally, we investigate the impact of halo mass dependence on these properties. Our conclusions are summarized below:

(i) All hydro-dynamical simulations and SAMs show a decreasing HI content from the halo outskirts to the halo centre. G-X20RPS in SAMs and two hydro-simulations agree better with the observations. The decreasing trends also depend on the host halo mass. Both EAGLE and TNG can reproduce the rapid decreasing of HI content in cluster haloes, and the milder decreasing trends observed in lower mass haloes. G-X20RPS and L-A21 from SAMs also present this different decreasing trend in high-mass and low-mass haloes, but predict a milder decrease of HI fractions in cluster haloes.

(ii) Halo mass dependence is also found in sSFR as a function of cluster-centric radii. For lower-mass haloes, the average sSFR only decreases at the innermost bin. While in cluster haloes, the decline starts since R_{200} . The same is found for all hydro-dynamical simulations and SAMs, among which, EAGLE and TNG show a better agreement with data.

(iii) The declining trend of HI is stronger than the decline of sSFR and starts at larger radii, in both lower-mass and massive haloes. The

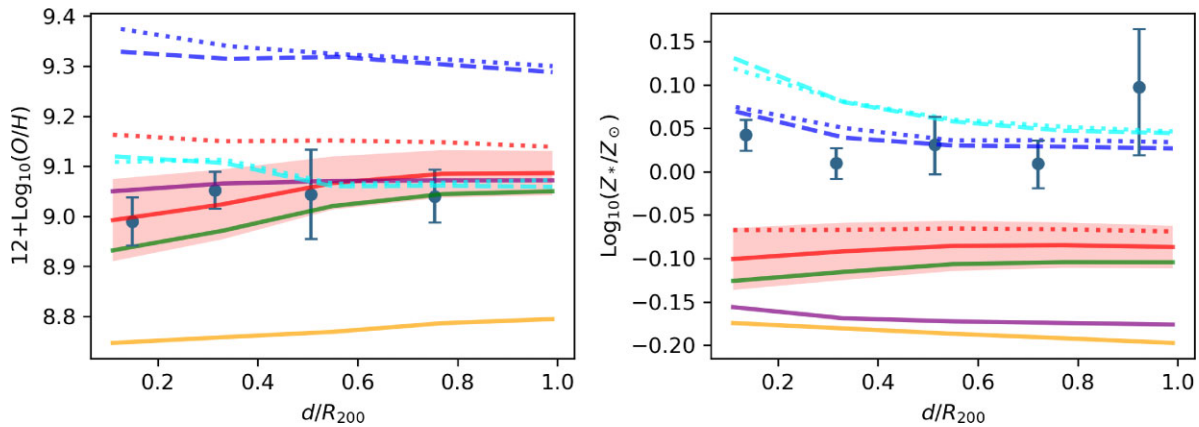


Figure 7. Radial projected distribution of oxygen abundance (left) of star-forming galaxies and stellar metallicity (right) of all galaxies, compared to observations (symbols). The projected distance is normalized by R_{200} . As in Fig. 4, solid curves show mean results, and dotted lines show the median values. Here, we use $Z_{\odot} = 0.02$. Note that the fifth bin (if shown) of the observation is unreliable due to its small sample size.

environmental effects have a more pronounced and earlier impact on the HI gas compared to their effect on star formation.

(iv) We find decreasing gas-phase metallicity for star-forming galaxies and increasing stellar-phase metallicity from the halo outskirts to the halo centre. The satellite galaxies would have increasing metallicity when suffering environmental effects. However, the selection of star-forming galaxies excludes metal-rich quenched satellite galaxies in halo centre and leads to a decrease in the gas-phase metallicity in halo centre.

Our findings highlight the crucial role of cold gas stripping in shaping the evolutionary trajectory of satellite galaxies within a dense environment. In the framework of GAEA, RPS stands out as the key reason for reproducing the drop of HI masses after galaxies fall into massive haloes. Also, the inclusion of sophisticated treatments of ram-pressure stripping for hot gas in L-A21 does improve the consistency with observational results with respect to L-H15. Both improvements are more obvious in high-mass haloes.

For EAGLE and TNG, the dramatic drop in the median HI mass, as well as the fact that half of satellite galaxies in halo central regions have no HI left, indicates much stronger suppression of the gaseous content for satellite galaxies in these hydro-dynamical simulations than that in semi-analytic models. Such a strong environmental effect might be, at least in part, attributed to artificial limited resolution or strong internal feedback that prevents gas from cooling and enhances gas stripping in hydro-dynamical simulations. Our work demonstrates that increasing the observational sample of HI measurements in massive haloes can provide useful constraints on the strength of environmental processes and ram-pressure stripping in particular.

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galaxies’. We are grateful to the referee for their many helpful suggestions.

DATA AVAILABILITY

An introduction to the GAEA code, as well as data file containing published model predictions, can be found at the GAEA webpage <https://sites.google.com/inaf.it/gaea/>. The data underlying this article will be shared on reasonable request to the corresponding author.

REFERENCES

- Abazajian K. N. et al., 2009, *ApJS*, 182, 543
 Ayromlou M., Kauffmann G., Yates R. M., Nelson D., White S. D. M., 2021b, *MNRAS*, 505, 492 (L–A21)
 Ayromlou M., Nelson D., Yates R. M., Kauffmann G., Renneby M., White S. D. M., 2021a, *MNRAS*, 502, 1051
 Ayromlou M., Nelson D., Yates R. M., Kauffmann G., White S. D. M., 2019, *MNRAS*, 487, 4313
 Bahé Y. M., McCarthy I. G., 2015, *MNRAS*, 447, 969
 Bahé Y. M., Schaye J., Crain R. A., McCarthy I. G., Bower R. G., Theuns T., McGee S. L., Trayford J. W., 2017, *MNRAS*, 464, 508
 Baldry I. K. et al., 2012, *MNRAS*, 421, 621
 Balogh M. L., Morris S. L., 2000, *MNRAS*, 318, 703
 Beckmann R. S. et al., 2017, *MNRAS*, 472, 949
 Blitz L., Rosolowsky E., 2006, *ApJ*, 650, 933
 Bluck A. F. L. et al., 2016, *MNRAS*, 462, 2559
 Boselli A., Fossati M., Sun M., 2022, *A&AR*, 30, 3
 Bothwell M. S., Maiolino R., Kennicutt R., Cresci G., Mannucci F., Marconi A., Ciccone C., 2013, *MNRAS*, 433, 1425
 Brinchmann J., Charlot S., White S. D. M., Tremonti C., Kauffmann G., Heckman T., Brinkmann J., 2004, *MNRAS*, 351, 1151
 Brown T. et al., 2017, *MNRAS*, 466, 1275
 Brown T., Cortese L., Catinella B., Kilborn V., 2018, *MNRAS*, 473, 1868
 Brüggén M., De Lucia G., 2008, *MNRAS*, 383, 1336
 Catinella B. et al., 2018, *MNRAS*, 476, 875
 Chabrier G., 2003, *PASP*, 115, 763
 Cora S. A. et al., 2018, *MNRAS*, 479, 2
 Cortese L., Catinella B., Boissier S., Boselli A., Heinis S., 2011, *MNRAS*, 415, 1797
 Cox T. J., Jonsson P., Primack J. R., Somerville R. S., 2006, *MNRAS*, 373, 1013
 Crain R. A. et al., 2015, *MNRAS*, 450, 1937
 Crain R. A. et al., 2017, *MNRAS*, 464, 4204
 Croton D. J. et al., 2006, *MNRAS*, 365, 11

- Davis M., Efstathiou G., Frenk C. S., White S. D. M., 1985, *ApJ*, 292, 371
- De Lucia G., Blaizot J., 2007, *MNRAS*, 375, 2
- De Lucia G., Tornatore L., Frenk C. S., Helmi A., Navarro J. F., White S. D. M., 2014, *MNRAS*, 445, 970
- De Lucia G., Weinmann S., Poggianti B. M., Aragón-Salamanca A., Zaritsky D., 2012, *MNRAS*, 423, 1277
- Diemer B. et al., 2018, *ApJS*, 238, 33
- Diemer B. et al., 2019, *MNRAS*, 487, 1529
- Dolag K., Borgani S., Murante G., Springel V., 2009, *MNRAS*, 399, 497
- Donnan C. T., Tojeiro R., Kraljic K., 2022, *Nat. Astron.*, 6, 599
- Donnari M. et al., 2019, *MNRAS*, 485, 4817
- Donnari M. et al., 2021a, *MNRAS*, 500, 4004
- Donnari M., Pillepich A., Nelson D., Marinacci F., Vogelsberger M., Hernquist L., 2021b, *MNRAS*, 506, 4760
- Dressler A., 1984, *ARA&A*, 22, 185
- Dressler A., Gunn J. E., 1983, *ApJ*, 270, 7
- Fabello S., Kauffmann G., Catinella B., Li C., Giovanelli R., Haynes M. P., 2012, *MNRAS*, 427, 2841
- Fontanot F. et al., 2021, *MNRAS*, 504, 4481
- Fu J., Kauffmann G., Li C., Guo Q., 2012, *MNRAS*, 424, 2701
- Gallazzi A., Charlot S., Brinchmann J., White S. D. M., Tremonti C. A., 2005, *MNRAS*, 362, 41
- Gao L., White S. D. M., Jenkins A., Stoehr F., Springel V., 2004, *MNRAS*, 355, 819
- Gavazzi G., O'Neil K., Boselli A., van Driel W., 2006, *A&A*, 449, 929
- Gnedin N. Y., Draine B. T., 2014, *ApJ*, 795, 37
- Gunn J. E., Gott J. Richard I., 1972, *ApJ*, 176, 1
- Henriques B. M. B., White S. D. M., Thomas P. A., Angulo R., Guo Q., Lemson G., Springel V., Overzier R., 2015, *MNRAS*, 451, 2663 (L–H15)
- Henriques B. M. B., Yates R. M., Fu J., Guo Q., Kauffmann G., Srisawat C., Thomas P. A., White S. D. M., 2020, *MNRAS*, 491, 5795 (L–H20)
- Hirschmann M., De Lucia G., Fontanot F., 2016, *MNRAS*, 461, 1760
- Hopkins P. F., Cox T. J., Hernquist L., Narayanan D., Hayward C. C., Murray N., 2013, *MNRAS*, 430, 1901
- Hough T. et al., 2023, *MNRAS*, 518, 2398
- Hu W. et al., 2019, *MNRAS*, 489, 1619
- Hu W., Cortese L., Staveley-Smith L., Catinella B., Chauhan G., Lagos C. D. P., Oosterloo T., Chen X., 2021, *MNRAS*, 507, 5580
- Hughes T. M., Cortese L., Boselli A., Gavazzi G., Davies J. I., 2013, *A&A*, 550, A115
- Jones M. G., Haynes M. P., Giovanelli R., Moorman C., 2018, *MNRAS*, 477, 2
- Kauffmann G. et al., 2003, *MNRAS*, 341, 33
- Knobel C. et al., 2013, *ApJ*, 769, 24
- Krumholz M. R., McKee C. F., Tumlinson J., 2009, *ApJ*, 693, 216
- Kukstas E. et al., 2023, *MNRAS*, 518, 4782
- Kulier A. et al., 2023, *ApJ*, 954, 177
- Larson R. B., Tinsley B. M., Caldwell C. N., 1980, *ApJ*, 237, 692
- Ma X., Hopkins P. F., Faucher-Giguère C.-A., Zolman N., Muratov A. L., Kereš D., Quataert E., 2016, *MNRAS*, 456, 2140
- Marasco A., Crain R. A., Schaye J., Bahé Y. M., van der Hulst T., Theuns T., Bower R. G., 2016, *MNRAS*, 461, 2630
- Marinacci F. et al., 2018, *MNRAS*, 480, 5113
- McAlpine S., Helly J. C., Schaller M., et al., 2016, *A&C*, 15, 72
- McKee C. F., Krumholz M. R., 2010, *ApJ*, 709, 308
- Moustakas J. et al., 2013, *ApJ*, 767, 50
- Naiman J. P. et al., 2018, *MNRAS*, 477, 1206
- Nelson D. et al., 2018, *MNRAS*, 475, 624
- Odekon M. C. et al., 2016, *ApJ*, 824, 110
- Peng Y.-j. et al., 2010, *ApJ*, 721, 193
- Peng Y., Maiolino R., Cochrane R., 2015, *Nature*, 521, 192
- Pillepich A. et al., 2018, *MNRAS*, 473, 4077
- Planck Collaboration XIII, 2016, *A&A*, 594, A13
- Planck Collaboration XVI, 2014, *A&A*, 571, A16
- Rahmati A., Pawlik A. H., Raičević M., Schaye J., 2013, *MNRAS*, 430, 2427
- Reiprich T. H., Basu K., Ettori S., Israel H., Lovisari L., Molendi S., Pointecouteau E., Roncarelli M., 2013, *Space Sci. Rev.*, 177, 195
- Saintonge A. et al., 2017, *ApJS*, 233, 22
- Sault R. J., Teuben P. J., Wright M. C. H., 1995, in Shaw R. A., Payne H. E., Hayes J. J. E. eds, *ASP Conf. Ser. Vol. 77, Astronomical Data Analysis Software and Systems IV*. Astron. Soc. Pac., San Francisco, p. 433
- Schaye J. et al., 2015, *MNRAS*, 446, 521
- Shi J. et al., 2022, *ApJ*, 927, 189
- Somerville R. S., Davé R., 2015, *ARA&A*, 53, 51
- Springel V. et al., 2005, *Nature*, 435, 629
- Springel V. et al., 2018, *MNRAS*, 475, 676
- Springel V., 2005, *MNRAS*, 364, 1105
- Springel V., 2010, *MNRAS*, 401, 791
- Springel V., White S. D. M., Tormen G., Kauffmann G., 2001, *MNRAS*, 328, 726
- Steinhauser D., Schindler S., Springel V., 2016, *A&A*, 591, A51
- Stevens A. R. H. et al., 2019, *MNRAS*, 483, 5334
- Stevens A. R. H., Brown T., 2017, *MNRAS*, 471, 447
- Tonnesen S., Bryan G. L., 2021, *ApJ*, 911, 68
- Torrey P. et al., 2019, *MNRAS*, 484, 5587
- van den Bosch F. C., Aquino D., Yang X., Mo H. J., Pasquali A., McIntosh D. H., Weinmann S. M., Kang X., 2008, *MNRAS*, 387, 79
- van den Bosch F. C., Jiang F., Campbell D., Behroozi P., 2016, *MNRAS*, 455, 158
- Vollmer B., Wong O. I., Braine J., Chung A., Kenney J. D. P., 2012, *A&A*, 543, A33
- Wang E. et al., 2018a, *ApJ*, 860, 102
- Wang E., Wang H., Mo H., van den Bosch F. C., Lim S. H., Wang L., Yang X., Chen S., 2018b, *ApJ*, 864, 51
- Wang K., Wang X., Chen Y., 2023, *ApJ*, 951, 66
- Warren M. S., Abazajian K., Holz D. E., Teodoro L., 2006, *ApJ*, 646, 881
- Wetzell A. R., Tinker J. L., Conroy C., 2012, *MNRAS*, 424, 232
- Woo J. et al., 2013, *MNRAS*, 428, 3306
- Xie L., De Lucia G., Hirschmann M., Fontanot F., 2020, *MNRAS*, 498, 4327
- Xie L., De Lucia G., Hirschmann M., Fontanot F., Zoldan A., 2017, *MNRAS*, 469, 968
- Yang X., Mo H. J., van den Bosch F. C., Zhang Y., Han J., 2012, *ApJ*, 752, 41
- Yates R. M., Henriques B. M. B., Fu J., Kauffmann G., Thomas P. A., Guo Q., White S. D. M., Schady P., 2021, *MNRAS*, 503, 4474
- Yates R. M., Henriques B., Thomas P. A., Kauffmann G., Johansson J., White S. D. M., 2013, *MNRAS*, 435, 3500
- Zinger E. et al., 2020, *MNRAS*, 499, 768
- Zinger E., Dekel A., Kravtsov A. V., Nagai D., 2018, *MNRAS*, 475, 3654
- Zu Y., 2020, *MNRAS*, 496, 111
- Zwaan M. A., Meyer M. J., Staveley-Smith L., Webster R. L., 2005, *MNRAS*, 359, L30

APPENDIX A: PROJECTION EFFECT

Here, we present the projection effect under different projection depths. We present its effect on the properties we discussed in this paper using G-X20RPS as an example, as shown in Fig. A1. The differences among project depths are negligible in all cases.

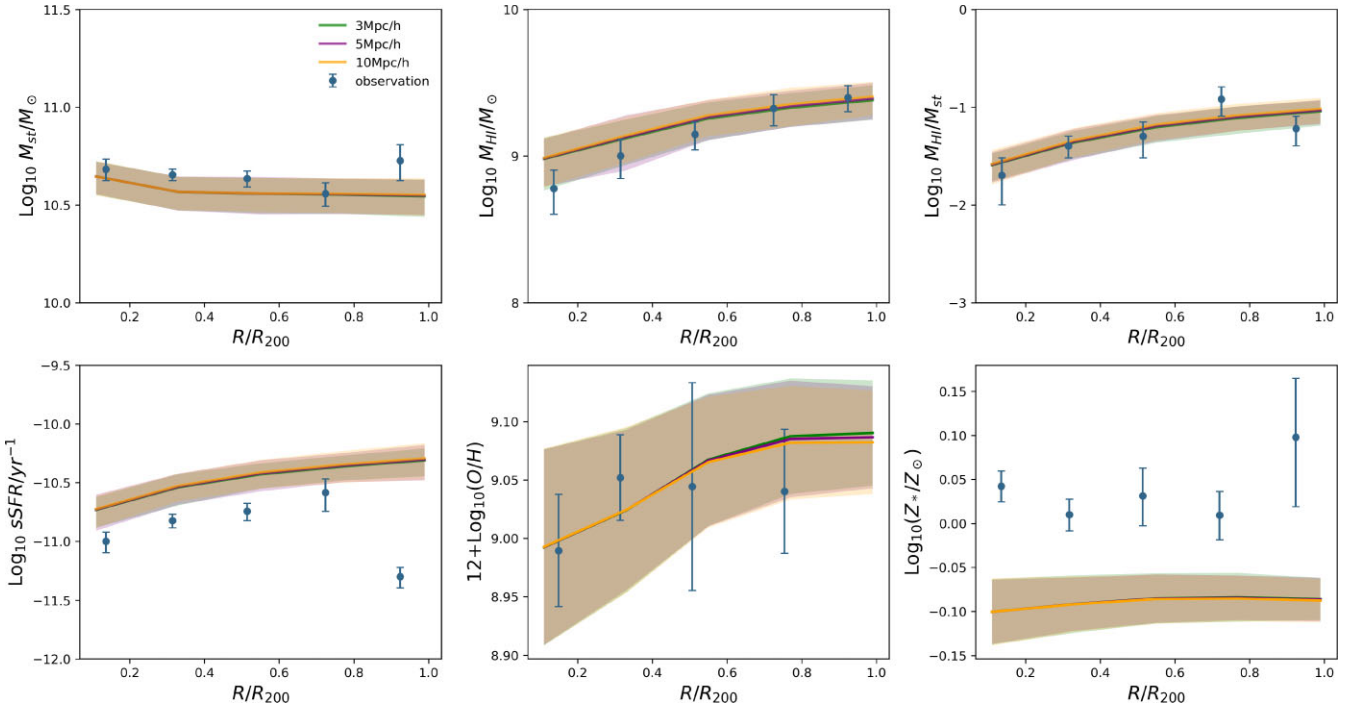


Figure A1. Radial properties of stellar mass, H I mass, H I mass fraction, sSFR, the gas-phase metallicity of star-forming galaxies and stellar metallicity under different projection depth 3 Mpc h^{-1} , 5 Mpc h^{-1} , and 10 Mpc h^{-1} .

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