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Images of Betelgeuse with VLTI/MATISSE across the Great Dimming

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ABSTRACT

From Nov. 2019 to May 2020, the red supergiant star Betelgeuse experienced an unprecedented drop of brightness in the visible domain called the great dimming event. Large atmospheric dust clouds and large photospheric convective features are suspected to be responsible for it. To better understand the dimming event, we used mid-infrared long-baseline spectro-interferometric measurements of Betelgeuse taken with the VLTI/MATISSE instrument before (Dec. 2018), during (Feb. 2020), and after (Dec. 2020) the GDE. We present data in the 3.98 to 4.15 μm range to cover SiO spectral features molecules as well as adjacent continuum. We have employed geometrical models, image reconstruction, as well as radiative transfer models to monitor the spatial distribution of SiO over the stellar surface. We find a strongly in-homogeneous spatial distribution of SiO that appears to be looking very different between our observing epochs, indicative of a vigorous activity in the stellar atmosphere. The contrast of our images is small in the pseudo-continuum for all epochs, implying that our MATISSE observations support both cold spot and dust cloud model.

Key words: (stars:) massive - (stars:) imaging - (stars:) individual - (stars:) variables: general - techniques: interferometric - techniques: image processing

1 INTRODUCTION

The red supergiant star Betelgeuse – Betelgeuse (α Ori-onis) is an O-rich star (Perrin et al. 2007) with a sun-like metallicity (Meynet et al. 2013) and a mass-loss rate of $1.2 \times 10^{-6} M_{\odot} \text{yr}^{-1}$ (Le Bertre et al. 2012). This red

supergiant (RSG) of spectral type classified as M2I by Levesque et al. (2005) has a mass estimated between 16.5 and 19 M_{\odot} by Joyce et al. (2020). The same authors give a distance of 168^{+28}_{-15} pc using seismic analysis, while Harper et al. (2017) give 222^{+48}_{-34} pc using new radio data taken with e-MERLIN, and older ALMA and VLA data

combined with the Hipparcos catalogue (van Leeuwen 2007). Its effective temperature T_{eff} is estimated to range from 3650 ± 25 K, using spectro-photometric observations (Levesque et al. 2005), to 3690 ± 54 K (Ohnaka 2014), using VLTI/AMBER spectro-interferometric observations. It is also a semi-regular variable star (SRc) with two main periods of ≈ 400 days and ≈ 2000 days identified by Gray (2008).

The great dimming scenario – From Nov. 2019 up to May. 2020, Betelgeuse went through an unexpected V-magnitude increase up to an historical value of 1.614 ± 0.008 (Guinan et al. 2020) reached between the 7th and the 13th of Feb. 2020. In this paper, we refer this 6 months event as the "great dimming event" (GDE). The proposed steps of the scenario for the GDE are:

Step I: Formation of a large plasma outflow [Feb. 2018–Jan. 2019] – Deep successive shocks generate a huge convective outflow gradually bringing the material to the surface. Kravchenko et al. (2021) detect two shocks located beneath the stellar photosphere, the stronger one in Feb. 2018 and the weaker one in Jan. 2019, using high resolution spectroscopic observations ($R=86000$) between 0.38 and $0.90 \mu\text{m}$ with HERMES (Raskin et al. 2011) at La Palma. The second shock, increasing the effect of the first one, contributes to the progressive formation of a plasma outflow at the surface of the photosphere;

Step II: Hot spot formation [Sep.–Nov. 2019] – The material growing up at the surface of the photosphere creates a hot bubble. Dupree et al. (2020) show an increase of the UV signal in the lines and in the continuum from Sept. to Nov. 2019 using Hubble observations. They interpret it as the sign of a hot, dense, luminous and high-temperature structure located in the southern hemisphere of the star between the photosphere and the chromosphere;

Step III: Cold spot and dust formation [Nov. 2019–Feb. 2020] – The new material detached from the photosphere forms progressively a gas cloud above the surface. Because of the missing material, a colder area appears on the photosphere seen as a dark spot in the line of sight and it allows the formation of dust in part of the cloud located above it. Spectro-photometric observations between 0.4 and $0.68 \mu\text{m}$, made by Levesque & Massey (2020) to estimate the effective temperature gradient induced by the GDE. They show no temperature variation between March 2004 and Feb. 2020, while Harper et al. (2020b) show a decrease of 125 K from Sept. 2019 to Feb. 2020 between 0.719 and $1.024 \mu\text{m}$. They explain these two apparently inconsistent results because of the presence of an opaque area made of cooler material covering at least more than half of the stellar surface. In addition, the presence of a dark spot at the surface of the star coupled with extinction of dust located above the surface agree with the VLTI/SPHERE visible observations made by Montargès et al. (2021) and VLTI/MATISSE N-band observations made by Cannon et al. (2023). It is also in good agreement with observations in polarized visible light of Cotton et al. (2020). They show a decrease of the polarization degree between Oct. 2019 and Feb. 2020, and explain it as a decrease of asymmetric illumination of the stellar surface and/or changes in the dust distribution.

Table 1. Observation logbook.

Date	Timespan	Cals.	AT-config.	Seeing ["]	τ_0 [ms]	Obs. ^a
2018-12-						
03	04:35–07:55	1 ^b , 2 ^c	A0B2J2C1	0.9	2.9	12
05	03:49–04:14	2	A0B2J2C1	0.7	1.5	6
08	01:45–07:29	1, 2	A0B2D0C1	0.6	3.5	24
09	03:30–07:53	1, 2	K0B2D0J3	0.4	4.8	17
11	03:27–07:52	1, 2	K0G2D0J3	0.6	8.7	24
12	04:30–06:16	1	A0G1D0J3	0.6	5.7	8
14	04:04–07:23	1	A0G1J2J3	0.6	5.6	8
15	04:02–05:50	1, 2	A0G1J2K0	0.5	5.8	12
2020-02-						
08	00:01–02:01	1	A0B2D0C1	0.6	5.9	12
19	00:19–01:37	2	K0B2D0J3	0.9	3.5	6
2020-12-						
14	04:10–07:21	1, 2	A0G1J2J3	0.5	10.1	24
18	04:01–05:17	1	K0G2D0J3	0.6	5.6	12
27	01:56–03:31	2	A0B2D0C1	0.8	5.0	18

^a Number of observing blocks actually used for the reduction and calibration of the data, ^b ε Lep (K4III, $\theta_L=5.7 \pm 0.5$ mas), ^c α CMa (A1V, $\theta_L=5.8 \pm 0.5$ mas).

For our study, we use L-band data taken with the VLTI/MATISSE instrument (Lopez et al. 2022) in medium spectral resolution ($R \approx 500$) before, during, and after the GDE to probe and monitor the geometry of the close circumstellar environment of Betelgeuse. The present letter is organized as follows: i) in Section 2 we present new VLTI/MATISSE observations of Betelgeuse in L-band; ii) in Section 3 we introduce the spectral analysis and the uniform disk fitting; iii) in Section 4 we show and compare the first reconstructed images of the SiO environment of Betelgeuse for the three epochs; and iv) in Section 5 we present our interpretations and conclusions.

2 OBSERVATION AND DATA REDUCTION

The three observing epochs of Dec. 2018, Feb. 2020, and Dec. 2020 are located before, during, and after the GDE of Betelgeuse (The N band data of the GDE already shown in Cannon et al. (2023)). All the MATISSE observations are close to a light curve photometric minimum of variability.

We reduced and calibrated the observation data using the MATISSE data reduction pipeline and additional python tools provided by the MATISSE consortium¹. The pipeline provides calibrated oifits files of the science target comprising measurements of the spectral distributions of flux, squared visibility, differential and closure phases, in the LM- and N-bands of the mid-infrared. Finally, we reject each night contaminated by adverse instrumental or weather effects, not already identified by the MATISSE pipeline. In Table 1 we list only the remaining nights after our sorting.

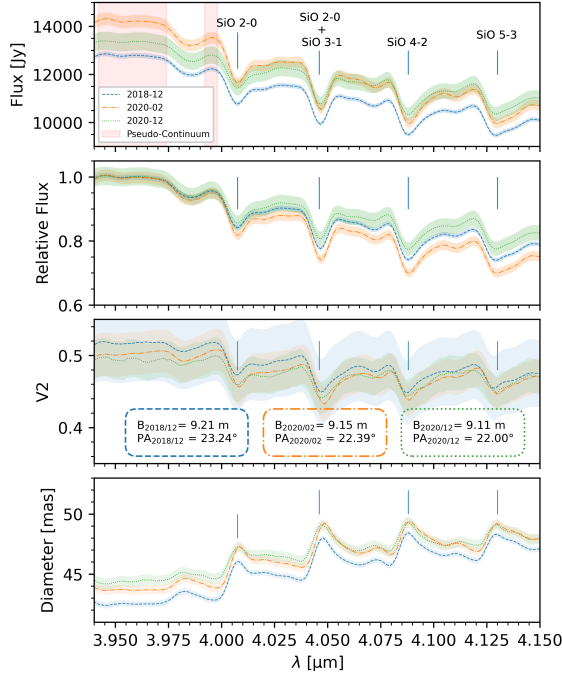


Figure 1. Top panel: VLTI/MATISSE absolute spectra for the three epochs with the identification of the main features. The filled area close to the data point corresponds to the error bars associated to the given quantities. The red area represent the pseudo-continuum range used in this work. Second panel: Relative flux with respect to the continuum for the three epochs. Third panel: visibility squared plotted versus wavelengths for the various epochs of the observations. Bottom panel: fitted uniform disk diameter versus wavelengths for the various epochs.

3 SPECTRAL ANALYSIS AND UNIFORM DISK FITTING

Figure 1 top panel shows a comparison between the three epochs of the Betelgeuse spectrum in the $3.98\text{--}4.15\ \mu\text{m}$ range. The SiO lines are weak, $\approx 20\%$ of the pseudo-continuum, and are in absorption. We observe changes in the spectrum between the 3 epochs. The pseudo-continuum level defined in this work between $3.942\text{--}3.974\ \mu\text{m}$ and $3.992\text{--}3.998\ \mu\text{m}$ seems to be brighter than pre-dimming by $\approx 10\%$ during the GDE, and by $\approx 5\%$ post-dimming. In the second panel of the figure we remark that the normalized flux to the continuum show SiO absorption features located at $4.01\ \mu\text{m}$, $4.05\ \mu\text{m}$, $4.09\ \mu\text{m}$, and $4.13\ \mu\text{m}$ (Aringer et al. 1997) which are slightly deeper during the Great Dimming (Feb. 2020) than before and after the event (Dec. 2018 and 2020). In the third panel we show the squared visibility for similar baselines and position angles for the three epochs. We observe small differences between the various epochs, these are within the measurement uncertainties, therefore it is hard to conclude a trend. In the bottom panel of Figure 1, we show the fitted equivalent uniform disk diameter obtained. The diameter’s variation for a given epoch reflects the presence

Table 2. Photometric quantities of the reconstructed maps. δI is the intensity contrast as defined by Wedemeyer-Böhm & Rouppe van der Voort (2009) estimated with an error bar of 6%, and MAD is the median absolute deviation of the intensity. The squared pixel size is 0.78 mas.

Date	Band	δI	Median [Jy/pix]	MAD [Jy/pix]
2018-12	Pseudo-cont.	0.12	4.4	0.3
2020-02	Pseudo-cont.	0.14	5.1	0.5
2020-12	Pseudo-cont.	0.12	4.7	0.3
2018-12	SiO (2–0)	0.14	3.4	0.3
2020-02	SiO (2–0)	0.16	3.8	0.4
2020-12	SiO (2–0)	0.11	3.6	0.3

of one or several layers of material above the stellar photosphere. The equivalent visibility-fit diameter appears smaller in Dec. 2018, while the star appears larger during and after the dimming event. This change has not to be interpreted as an actual change of the star’s photospheric diameter, but rather a change of brightness of the atmosphere layers (see also Fig. 2). In addition, the error bars associated to the angular diameter look smaller regarding the errors of the squared visibilities shown above, due to the fact that the first-lobe uniform disk fitting is constrained by dozens of squared visibilities measurements (and not only the values shown in the third panel of Figure 1).

4 IMAGES

We use the image reconstruction software IRBIS (Hofmann et al. 2014) to produce L-band maps at the three epochs of observation in the combined spectral ranges $3.942\text{--}3.974\ \mu\text{m}$ and $3.992\text{--}3.998\ \mu\text{m}$ (pseudo-continuum), and in the spectral range $4.004\text{--}4.012\ \mu\text{m}$ (SiO 2–0 transition). After a cautious scan of the imaging parameter space, we found the following optimal parameters for the reconstruction: i) initial image: an uniform disk, ii) prior: maximum entropy, and iii) hyperparameter: $\mu = 10$. We show the obtained images in Figure 2. Since the Feb. 2020 observation run has a sparse (u,v)-plane coverage, in order to compare the three epochs images together, we convolved them with the Feb. 2020 equivalent circular interferometric beam of an angular diameter of 4 mas.

In order to make a quantitative analysis between the reconstructed maps obtained at the 3 epochs, we compute the values of the photometric median, the median absolute deviation (MAD), and the intensity contrast as defined by Wedemeyer-Böhm & Rouppe van der Voort (2009). The values reported in Table 2 are computed for intensities higher than 70% of the maximum (Chiavassa et al. 2022) to exclude the region of the disk dominated by the limb darkening effect. The corresponding size measures about 42 mas, shown as dashed white circles in the reconstructed maps of Figure 2. In addition, we use the equations developed in Appendix B to estimate the optical depth map of SiO, shown in the center-right panels of Figure 2. This allows us to monitor the SiO gas distribution across the three epochs.

5 DISCUSSION

We comment here the $4\ \mu\text{m}$ maps reconstructed from the VLTI/MATISSE instrument obtained before, during, and af-

¹ mat_tools v.0.1 freely available on Gitlab <https://gitlab.oica.eu/MATISSE/tools/-/tree/master>

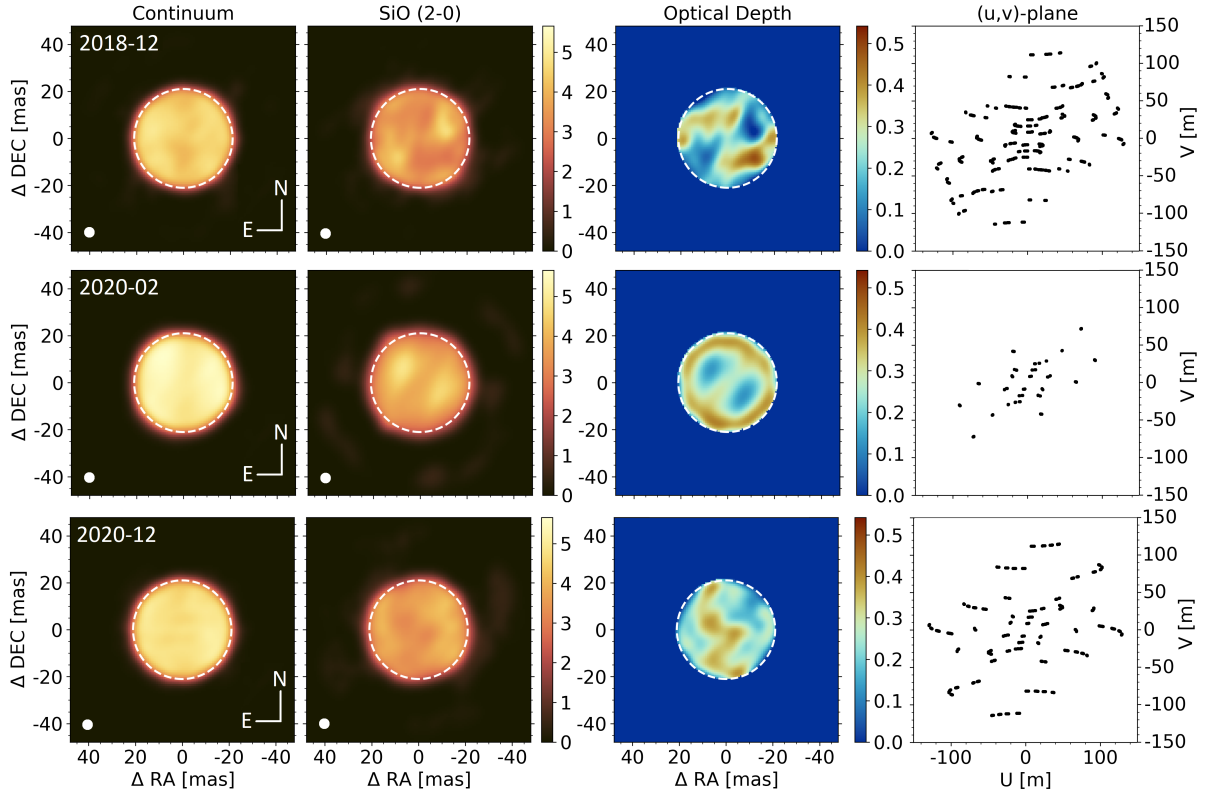


Figure 2. Each rows correspond to a given epoch precised in the top-left corner of the first panel. From left to right: 1) Reconstructed map in the pseudo-continuum, 2) in the SiO (2–0) absorption band (both scaled in Jy/pix with a squared pixel size of 0.78 mas), 3) deduced map of the optical depth for the SiO (2–0), 4) (u,v)-plane coverage of the used measurements. The dashed white circles on the maps enclose intensities higher than 70% of the maximum (42 mas in diameter). The small white disks at the bottom left corners of the reconstructed maps corresponds to an equivalent circular interferometric beam used for the convolution of all the reconstructed images (≈ 4 mas in diameter). The angular size has been estimated using the Feb. 2020 observations which has the poorest (u,v)-plane coverage among the three epochs.

ter the GDE in the SiO (2–0) absorption band and in the continuum.

5.1 Wavelength and Time variations of the angular diameter

The uniform disk angular diameter, obtained in section 3, exhibit chromatic variations. These variations give hints to the geometry of the SiO dust seed formation region of Betelgeuse. Indeed, for example in Dec. 2018, while the apparent uniform disk diameter is ≈ 42.5 mas in the pseudo-continuum, it is ≈ 46.0 mas for the SiO 2–0 transition and ≈ 48 mas for the remaining 3 transitions (SiO 2–0 + SiO 3–1, SiO 4–2, and SiO 5–3), meaning that, qualitatively, the SiO layer extends 2 mas above the stellar photosphere, and spans 2 mas thick. Of course, such a qualitative analysis should be supplemented by more quantitative ones, using state-of-the-art molecular models of Betelgeuse atmosphere to constrain the vertical structure of the SiO distribution, but it is out of the scope of this letter.

The UD angular diameter shows also time variations: figure 1 indicates a global increase of 3% of the visibility-fitted uniform disk angular diameter between Dec. 2018 and Feb. 2020. From 42.5 mas to 43.8 mas in the continuum and 46.0 mas to 47.4 mas in the SiO (2–0) transition band, this small increase is about the same order of magnitude as the 1% increase reported in K-band (Montargès et al. 2021).

However, as explained by Dharmawardena et al. (2020) and Montargès et al. (2021), such a small variation of the angular diameter is not sufficient to explain the GDE.

5.2 Time variation of the pseudo-continuum

The intensity contrast gives important quantitative information about the distribution of the surface brightness intensity of the star. Its uncertainty value of 6% has been estimated by measuring the variation of the intensity contrast on reconstructed images from simulated interferometric data of a uniform disk (see images in figure A1, and more details in section A). A temporal intensity contrast comparison between different images allows us to probe the surface and atmospheric activities between the different periods. In Table 2 we report an increase of the intensity contrast over the surface in the pseudo-continuum images around $4 \mu\text{m}$ of $\sim 17\%$ between Dec. 2018 (before the GDE) and Feb. 2020 (during the GDE). Coupling those quantities with a slight increase of the median flux, this transcribe the appearance of a global brighter surface during the GDE which might be associated to recently dust formed in the line of sight. Figure 1 also shows a global increase of the infrared emission in the continuum of $\sim 13\%$ which is in agreement with the presence of dusty material recently formed close to the star as proposed by Montargès et al. (2021). The results of

Harper et al. (2020a), who do not find any infrared excess around $25\ \mu\text{m}$, are not necessarily in disagreement with our results and might be complementary to constrain the nature of the chemical composition of the newly formed dust in the environment of the star during the GDE (but this last point is out of the scope of the current work).

5.3 Time variation of the SiO asymmetries

Since the quality and the robustness of the image reconstruction strongly depend on the completeness of the (u,v) -plane coverage, we need to be careful about our interpretation on the structures shown in the images, especially when the (u,v) -plane is poor like the Feb. 2020 one as shown in Figure 2. Indeed, the presence of artifacts in the images may impair our interpretation. They are flux whose owns a poorly constraint shape due to the holes in the (u,v) -plane coverage.

Assuming that dust seeds precursors such as SiO molecule can be used as tracers for dust formation region, therefore asymmetries in the gas distribution can also be a way to probe asymmetries in the distribution of recently formed dust. In Dec. 2018 and Dec. 2020, we observe strong and different asymmetries in the SiO (2–0) gas distribution. This is in agreement with some episodic and non-homogeneous ejection of mass expected for a RSG (Toalá & Arthur 2011). Regarding the Dec. 2020 images, we cannot exclude artifacts contamination induced by poor coverage of the (u,v) -plane. Therefore we cannot make any strong conclusions about changes in the gas distribution during this epoch.

In the SiO (2–0) band, the intensity contrast increases during the GDE by 14% with an increase of the median flux of 10% with respect to Dec. 2018. Therefore, it seems that during the GDE we observe brighter structures in the line of sight. Next, we observe in Dec. 2020 a decrease at about 50% of the intensity contrast with a decrease of the median flux of 7% regarding Feb. 2020. The environment seems to be smoother compared to Feb. 2020 with a median value getting closer from the Dec. 2018 values.

Looking at the optical depth map of Dec. 2018 and Dec. 2020, that have a denser (u,v) -plane coverage than Feb. 2020, we can identify opacity variations up to a factor 5 and 3, far more than the estimated 6% error (section 5.2), respectively. The SiO (2–0) opacity depth map shows therefore strong temporal variations within 2 years, indicative of vigorous changes in the star’s environment in this time span.

5.4 Dust Clump and Cold Spot model

In addition we checked the robustness of the cold spot and dust clump model used in Montargès et al. (2021) to explain the GDE with respect to our Feb. 2020 VLT/MATISSE data. Both models are derived from radiative transfer code and own therefore wavelength-dependent structures. Here, we use the exact same models with the same parameters as in Montargès et al. (2021, the reader can refer to that paper to get the models parameters values), just in a different wavelength domain. Thus, the dust cloud and the cold spot presented in this paper as seen by MATISSE in L-band (4000 nm), may not exhibit the same dynamic range as in Montargès et al. (2021) in V-band ($\approx 650\ \text{nm}$).

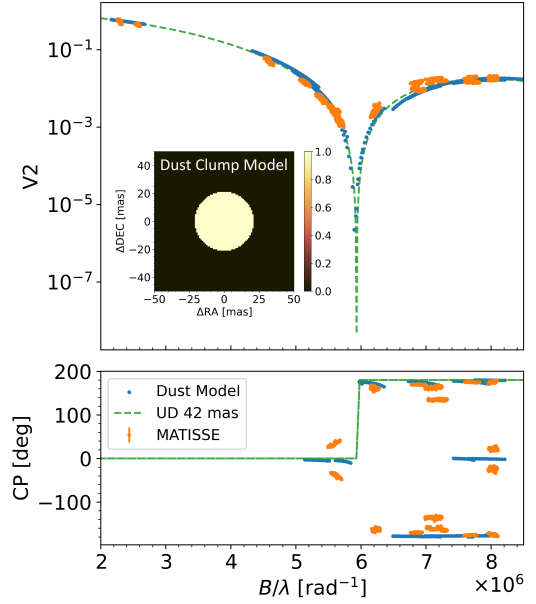


Figure 3. Upper panel: Visibilities squared in the continuum of the MATISSE data from Feb. 2020 (orange), dust patch model (blue) with its image in the left corner, and an uniform disk of angular diameter 42 mas (green). Lower panel: representation of their respective closure phase with respect to the spatial frequency of the longest baseline in the baselines triplet.

Indeed, in the top panel of Figure 3 showing the dust clump model, we can see that the dynamic range of the model image (inset) used to compute the L-band visibilities (blue dots) is insufficient to distinguish the models from an uniform disk model (green dashed line). The observed data (orange dots) exhibit small deviation from both the Uniform disk model and the dust clump model. Concerning the closure phase, the dust clump model exhibits a slight deviation from zero or 180° closure phase (i.e. deviation from the centro-symmetric geometry), but not sufficient enough to reproduce the MATISSE data. The same conclusions can be drawn looking at Figure C1 for the cold spot model. The visibilities are also consistent with the MATISSE data but the closure phase does not seems to show any deviation from zero or 180° closure phase. Thus, both cold spot and dust clump models in L-band are consistent with the MATISSE visibilities but are not able to reproduce the closure phase.

6 CONCLUSION

Our observations suggest an optically thin molecular medium in the SiO (2–0) band absorption at the three epochs. They suggest also the presence of an infrared excess in the pseudo-continuum during the GDE, which has been interpreted as new hot dust formed. We also detect the presence of brighter structures in L-band appearing during the GDE both in the continuum and in the SiO bands. The MATISSE visibilities observations of Feb. 2020 are both compatible with the presence of a cold spot on the surface of the star and that of a cloud of dust formed in the line of sight, but these two models struggle to explain the observed closure phases. The Dec. 2020 observations suggests that Betelgeuse seems to be

returning to a gas and surface environment similar to the one observed in Dec. 2018 but with smoother structures maybe due to unusual amount of dust recently formed during the GDE in the line of sight. Opacity maps also show strong spatial variations in Dec. 2018 and Dec. 2020 epochs by respectively a factor of 5 and 3, evidencing a clear inhomogeneous formation of SiO in the stellar atmosphere, as well as vigorous changes in a timespan of 2 years. In addition to the opacity maps, the chromatic estimations of the equivalent uniform disks visibility-fitted angular diameters give also hints to constrain the geometry of the SiO dust seed formation region.

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Data availability: The data underlying this article are available in the OIdb digital repository with direct links to the data collections: [Dec. 2018](#), [Feb. 2020](#), and [Dec. 2020](#).

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APPENDIX A: IMAGE RECONSTRUCTION OF AN UNIFORM DISK WITH IRBIS

In order to identify possible artifacts in our images induced by the (u,v)-plane, we used the ASPRO2 software to simulate (u,v) observations with a similar (u,v) plane as for our observations, and we used the IRBIS software to reconstruct images based on these simulated data. These images (Figure A1) are made using the same reconstruction parameters as for the pictures in Figure 2. All the structures identified in Figure A1 which deviates from the uniform disk model are considered as artifacts. Except from the X-shaped structure observed during the Feb. 2020 epoch, the symmetric structures seen in our simulation do not seem to be present in Figure 2, hence we can assume that the variations observed in the reconstructed opacity map at the Dec. 2018 and Dec. 2020 epochs are real.

APPENDIX B: OPTICAL DEPTH MAP ESTIMATION

From these images we derive the spectral distribution of the optical depth at the SiO absorption bands $\tau_{\text{SiO}}(\lambda)$ assuming that:

- 1) the stellar photosphere is the dominant source of the pseudo-continuum radiation, denoted hereafter as $S_{\star}(\lambda)$;
- 2) the SiO material is responsible not only for the absorption but also for the emission of the light in the SiO absorption bands, denoted hereafter as $S_{\text{SiO}}(\lambda)$.

Under these assumptions, the radiative transfer equation becomes:

$$I(\lambda) = S_{\star}(\lambda) e^{-\tau_{\text{SiO}}(\lambda)} + S_{\text{SiO}}(\lambda) \left[1 - e^{-\tau_{\text{SiO}}(\lambda)} \right], \quad (\text{B1})$$

where $I(\lambda)$ is the spectral distribution of surface brightness. Since the absorption feature in the spectrum is weak (c.a. 20% of the pseudo-continuum), we assume that the SiO layer is optically thin, i.e. $\tau_{\text{SiO}}(\lambda_{\text{band}}) < 1$.

In order to get an estimate of the optical depth, we make the following additional simplifying assumptions:

- 3) the SiO contribution to the continuum wavelength range is negligible i.e. $\tau_{\text{SiO}}(\lambda_c) = 0$, hence:

$$I(\lambda_c) = S_{\star}(\lambda_c); \quad (\text{B2})$$

- 4) in front of the stellar disk, the SiO emission is negligible compared to the stellar emission in the absorption band, so that:

$$S_{\text{SiO}}(\lambda_{\text{band}}) \ll S_{\star}(\lambda_{\text{band}}), \text{ and} \quad (\text{B3})$$

$$S_{\text{SiO}}(\lambda_{\text{band}}) \left[1 - e^{-\tau_{\text{SiO}}(\lambda_{\text{band}})} \right] \approx 0, \quad (\text{B4})$$

hence:

$$I(\lambda_{\text{band}}) = S_{\star}(\lambda_{\text{band}}) e^{-\tau_{\text{SiO}}(\lambda_{\text{band}})}. \quad (\text{B5})$$

- 5) the stellar source function is the same in the absorption band as in the continuum as long as the continuum level does not significantly change between both spectral channels, so that:

$$S_{\star}(\lambda_c) = S_{\star}(\lambda_{\text{band}}). \quad (\text{B6})$$

Substituting these assumptions into equations B2 and B5 leads to a rough estimate of the optical depth at the wavelength range of the SiO (2–0) absorption band :

$$\tau_{\text{SiO}}(\lambda_{\text{band}}) = \log I(\lambda_c) - \log I(\lambda_{\text{band}}). \quad (\text{B7})$$

APPENDIX C: COLD SPOT MODEL

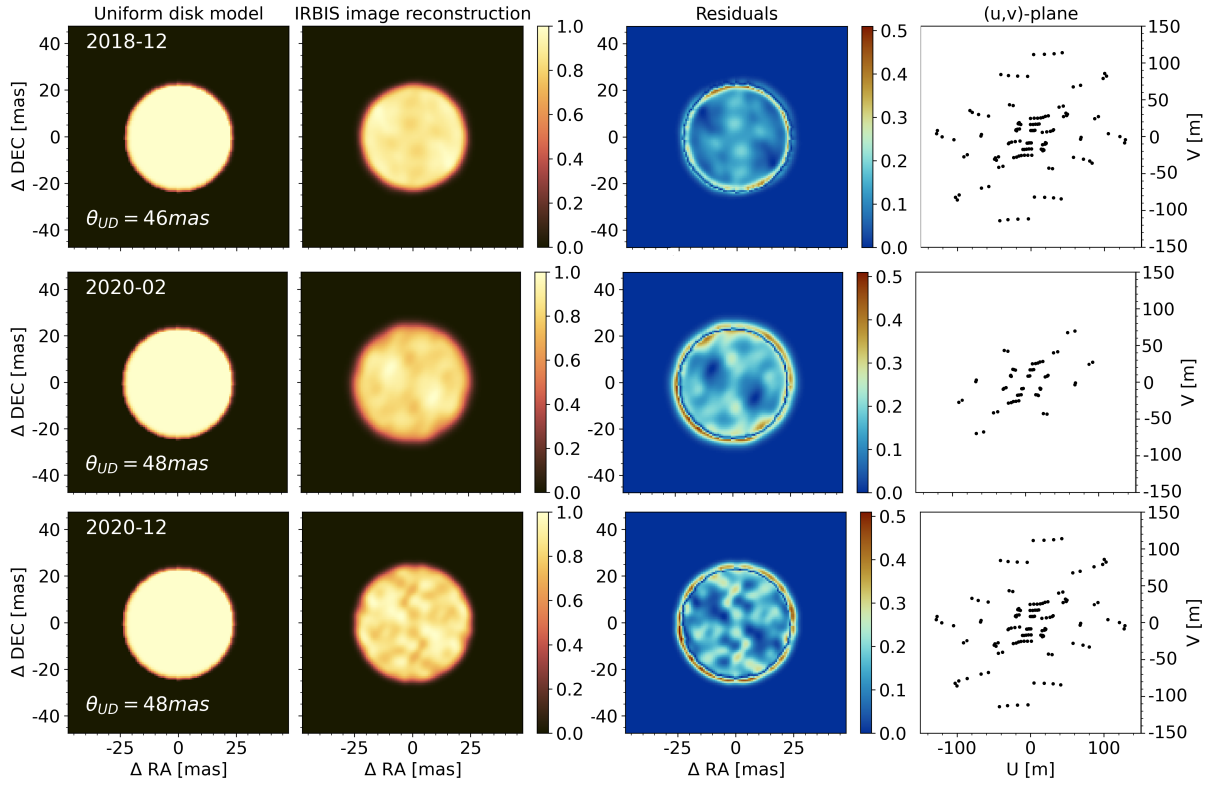


Figure A1. Each row correspond to a given epoch precised in the top-left corner of the first panel. From left to right: 1) model of the uniform disk convolved with the same interferometric beam as used in Figure 2, 2) reconstructed image using IRBIS of a simulated interferometric data of a uniform disk with an angular diameter determined using the values fitted on the visibility squared showed in Figure 1 for the SiO first overtone wavelength, 3) residuals between the model and the image reconstructed, 4) simulated (u,v)-plane coverage used for the image reconstruction of the simulated data.

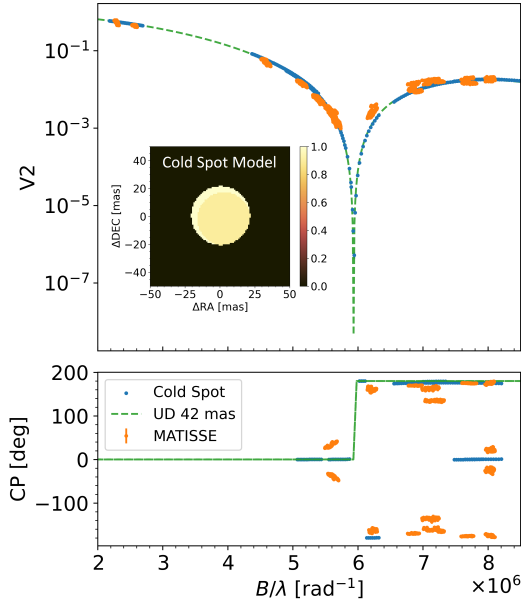


Figure C1. Upper panel: Visibilities Squared in the continuum of the MATISSE data of February 2020 (orange), cold spot model (blue) with its image in the left corner, and an uniform disk of angular diameter 42 mas (green). Lower panel: representation of their respective closure phase.