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Cosmology with cosmic web environments

II. Redshift-space auto and cross-power spectra

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ABSTRACT

Degeneracies among parameters of the cosmological model are known to drastically limit the information contained in the matter distribution. In the first paper of this series, we show that the cosmic web environments, namely the voids, walls, filaments and nodes, can be used as leverage to improve the real-space constraints on a set of six cosmological parameters, including the summed neutrino mass. Following up on these results, we propose to investigate the extent to which constraints can be obtained with environment-dependent power spectra in redshift space where the velocities add information to the standard two-point statistics by breaking the isotropy of the matter density field. A Fisher analysis based on a set of thousands of Quijote simulations allows us to conclude that a combination of power spectra computed in several cosmic web environments is able to break some degeneracies. Compared to the matter monopole and quadrupole information alone, the combination of environment-dependent spectra improves constraints on key parameters such as the matter density and the summed neutrino mass by up to a factor of 5.5. Additionally, while the information contained in the matter statistic quickly saturates at mildly non-linear scales in redshift space, a combination of power spectra from different environments appears to be a rich source of information that can be used to improve the constraints at all the studied scales from 0.1 to $0.5 h \text{ Mpc}^{-1}$ and suggests that further improvements could be attainable at even finer scales.

Key words. cosmology: theory – large-scale structure of Universe – cosmological parameters

1. Introduction

The spatial distribution of matter in the Universe is a powerful probe to constrain cosmological parameters (see e.g. Totsuji & Kihara 1969; Peacock et al. 2001; Eisenstein et al. 2005; Mandelbaum et al. 2013; Alam et al. 2017a; Hildebrandt et al. 2017; D'Amico et al. 2020; Colas et al. 2020; Ivanov et al. 2020), test the theory of gravity (Alam et al. 2017b, 2021; Jullo et al. 2019; Blake et al. 2020), and improve our understanding of dark matter and dark energy (Abbott et al. 2019; Drlica-Wagner et al. 2019). All these possibilities have led numerous collaborations over the past decades to initiate a race to map the largest number of galaxies in the sky with the highest possible accuracy. Stage IV of these galaxy redshift surveys – notably the Dark Energy Spectroscopic Instrument¹ (DESI Collaboration 2016), *Euclid*² (Laureijs et al. 2011) and the *Nancy Grace Roman* space telescope³ (Spergel et al. 2015) – promises an unprecedented amount of data, enabling, among other things, the most accurate estimate of cosmological parameters to date. However, these expected observational achievements must be accompanied by appropriate theoretical and numerical advances. Although widely used in most analyses because of its simplicity, it is

well-known that the power spectrum (or equivalently the two-point correlation function), for example, is not a sufficient summary statistic when it comes to non-Gaussian fields such as the late-time distribution of matter in the Universe. Another caveat that must be taken into account when using the matter power spectrum for cosmological analyses is the degeneracies among parameters of the Λ cold dark matter (Λ CDM) model, which have a similar impact over a wide range of scales, limiting the information one can extract from the model.

Over recent years, several strategies have been proposed to fight the degeneracies and go beyond simple power spectrum analyses in order to optimally exploit the matter distribution. These have included information coming from higher-order statistics (Hahn et al. 2020; Hahn & Villaescusa-Navarro 2021; Gualdi et al. 2021; Philcox & Ivanov 2022), velocities (Mueller et al. 2015; Kuruvilla & Aghanim 2021; Kuruvilla 2022), marked power spectra (Massara et al. 2021, 2022), wavelet-based statistics (Allys et al. 2020; Cheng et al. 2020; Cheng & Ménard 2021; Valogiannis & Dvorkin 2022a,b; Eickenberg et al. 2022; Wang & He 2022; Wang et al. 2022), split densities (Uhlemann et al. 2020; Paillas et al. 2021, 2023), partial or total cosmological environments (Kreisch et al. 2022; Bayer et al. 2021; Bonnaire et al. 2021; Woodfinden et al. 2022), or from the minimum spanning tree (Naidoo et al. 2020, 2022). While the recent literature abounds in analyses and forecasts designed to identify the optimal statistic in real space, only a few studies have

¹ <https://www.desi.lbl.gov/>

² <https://www.euclid-ec.org/>

³ <https://roman.gsfc.nasa.gov/>

addressed the problem in the space in which the data are actually collected, namely redshift space. It is nonetheless well-known that the additional information carried by velocities coming into play in this space is non-negligible and leads to considerable improvement of the cosmological constraints derived from the standard matter statistics.

The real-space results presented in Bonnaire et al. (2021) demonstrated the ability of simple two-point statistics derived in the separate cosmic web environments (voids, walls, filaments, and nodes) to break key degeneracies in the cosmological model, consequently enabling the improvement of constraints on all the parameters compared to those derived from the matter power spectrum. In the present study, we extended our analysis of the constraints obtained by the combination of power spectra computed in several environments when considering redshift-space distortions (RSDs). In particular, we investigated how the combination of environment-dependant spectra in this space can improve the cosmological constraints, and how these latter compare to the real-space case without correcting the distortions at all. To this end, we performed a Fisher forecast of the constraints using numerous simulations with varying cosmological parameters in order to numerically assess the information contained in our statistics on the cosmological parameters. The paper is organised such that Sect. 2 presents the Quijote suite of simulations and the methodology used to extract the cosmic web environments while Sect. 3 introduces the power-spectra estimation in redshift space and in the cosmic web environments. Section 4 then presents the forecast constraints in redshift space but also compares these with the real space findings reported in Bonnaire et al. (2021). Finally, Sect. 5 provides a discussion of the obtained results in view of the recent literature, gives perspectives for future steps towards the use of this kind of analysis in observational setups, and also contains a summary and our conclusions.

2. Data and methodology

2.1. The Quijote simulations

Our analysis relies on the Quijote⁴ (Villaescusa-Navarro et al. 2020) suite, which provides a set of 44 100 N -body simulations. With realisations of more than 7000 cosmological models, this large suite of simulations was precisely designed to perform statistical analyses and train machine learning algorithms. Initialised with the second-order Lagrangian perturbation theory (or Zel'dovich approximation in case of massive neutrino simulations), $N = 512^3$ dark matter particles (and 512^3 neutrinos in case there are) are evolved forward in time in an $L = 1 h \text{ Gpc}^{-1}$ size box from redshift $z = 127$ to $z = 0$. There are 15 000 simulations available at the fiducial cosmology, which is assumed to be a flat Λ CDM cosmology with parameters $\Omega_m = 0.3175$, $\Omega_b = 0.049$, $h = 0.6711$, $n_s = 0.9624$, $\sigma_8 = 0.834$, and $M_\nu = 0$, with M_ν being the summed neutrino mass. For each cosmological parameter, 500 simulations are then individually computed with a fixed increase and decrease, $d\Omega_m = \pm 0.010$, $d\Omega_b = \pm 0.002$, $dh = \pm 0.020$, $dn_s = \pm 0.020$, and $d\sigma_8 = \pm 0.015$. Because M_ν is a positive quantity, the Quijote suite provides four positive variations of this parameter with $M_\nu = \sum m_\nu = \{0.1, 0.2, 0.4\} \text{ eV}$.

⁴ <https://quijote-simulations.readthedocs.io/en/latest/>

2.2. Cosmic web classification

Among the many possible definitions of the cosmic web environments proposed in the literature (see Stoica et al. 2007; Aragon-Calvo et al. 2010; Cautun et al. 2013; Sousbie 2011; Tempel et al. 2014; Bonnaire et al. 2020, 2022, to name only a few), we resort to an implementation of the T-Web algorithm (Hahn et al. 2007; Forero-Romero et al. 2009). We summarise the general steps of the classification below. We first start by computing a smooth density field $\rho(\mathbf{x})$ based on the discrete set of particles by means of a B -spline interpolation of order four over an $N_g^3 = 360^3$ three-dimensional grid. The T-Web formulation relies on the tidal tensor $\mathbf{T}(\mathbf{x})$, the second derivative of the gravitational potential, itself computed from $\rho(\mathbf{x})$, and the Poisson equation. Each of the grid cells \mathbf{x} is classified as either void, wall, filament, or node depending on the eigenvalues of the tidal tensor. More precisely, a cell \mathbf{x} is in a void if the three eigenvalues are below λ_{th} ; in a wall if only two are below λ_{th} ; in a filament if only one is below λ_{th} ; and in a node if none are below this threshold. From the classification of the N_g^3 cells of the density field, we propagate the environments at the particle level, enabling the computation of four individual overdensity fields, δ_v , δ_w , δ_f , and δ_n , where the subscripts refer to void, wall, filament, and node, respectively. These fields are all linked to the full matter overdensity field δ_m by the linear combination $\delta_m = \sum_{\alpha \in \{v,w,f,n\}} f_\alpha \delta_\alpha$, where $f_\alpha = N_\alpha/N$ denotes the mass fraction⁵ of the environment α .

The classification scheme relies on three parameters. The first two are related to the coarseness of the smooth gravitational potential estimation, N_g^3 , the number of grid cells, and σ_N , the scale of the Gaussian with which the potential is being smoothed before performing the classification. The last parameter, λ_{th} , corresponds to the threshold used for the classification of the environments based on the amplitudes of the eigenvalues of the tidal tensor. In Bonnaire et al. (2021), we explored the impact of the parameters and ultimately set $N_g^3 = 360^3$ cells and a value of $\sigma_N = 2 h \text{ Mpc}^{-1}$ for the smoothing scale. This yields an effective smoothing of $3.4 h \text{ Mpc}^{-1}$, which corresponds to the physical size of the structures, such as galaxy clusters or filament widths (see Cautun et al. 2014). The eigenvalue threshold is set to $\lambda_{\text{th}} = 0.3$ in order to obtain a cosmic web that is consistent with the use of such a methodology in the literature and in which the voids start to percolate (see e.g. Hahn et al. 2007; Martizzi et al. 2019; Libeskind et al. 2017; Bonnaire et al. 2021). We refer the reader to the first paper of the series (Bonnaire et al. 2021) for details on the analysis and assessment of these choices of parameter values.

3. Redshift space

3.1. Redshift-space distortions

When dealing with observational data, the redshift is used as a measure of the distance. However, this quantity is the combination of the proper motion of the source due to its peculiar velocity, \mathbf{v} , and the expanding Universe. Considering \mathbf{x}_r , the position of a source in the comoving space, \mathbf{x}_s , its position in the redshift space, and $\hat{\mathbf{n}}$, a unit vector in the line of sight (LoS) direction,

⁵ As all the particles have the same mass in the simulation, the ratio between N_α , the number of particles in the environment α , and the total number of particles N gives the mass fraction.

we can obtain the mapping relation

$$\mathbf{x}_s = \mathbf{x}_r + \left(\frac{\mathbf{v} \cdot \hat{\mathbf{n}}}{aH} \right) \hat{\mathbf{n}}, \quad (1)$$

where a is the scale factor and H the Hubble constant. This shift distorts the spatial distribution of the tracers, and is known as redshift-space distortion (RSD), which leads to two dominant effects: The first is the finger-of-God (FoG) effect (Jackson 1972), which makes overdense clustered regions appear elongated in the LoS direction due to the high velocity of the sources, and the second, observed on larger scales, is a squashing effect of dense regions along the LoS known as the Kaiser effect (Kaiser 1987).

While Bonnaire et al. (2021) focused on the information content of power spectra in cosmic web environments computed in real space, with the current work we propose to investigate the information in redshift space, which is therefore expected to more closely resemble observations. To do so, we mimic the effect of RSDs in each individual simulation by displacing all particles (dark matter particles and neutrinos if any) according to Eq. (1) along one Cartesian axis of the box, and therefore assume that the plane-parallel approximation holds.

3.2. Power-spectra estimation in redshift-space

In a general manner, the cross-power spectrum is defined as the covariance of Fourier transformed overdensity fields δ_α and δ_β and is given by

$$P_{s,\alpha\beta}^\ell(k) = \frac{2\ell + 1}{2} \int_{-1}^1 P_{\alpha\beta}(k, \mu) \mathcal{L}_\ell(\mu) d\mu, \quad (2)$$

where $\mu = \mathbf{k} \cdot \hat{\mathbf{n}}/k$, which is the angle with the line of sight, $P_{\alpha\beta}(k, \mu)$ is the 2D power spectrum obtained by binning both in k and μ , and \mathcal{L}_ℓ are the Legendre polynomials. In real space, the isotropy of the density field implies that all the $\ell > 0$ terms cancel out, leaving the monopole alone to carry all the information. In redshift space, the peculiar velocities of particles induce a dependence of the power spectrum on the LoS, which leads to the breaking of the density field isotropy and consequently spreads the power over the multipoles ℓ . Throughout the present paper, we write $P_{\alpha\beta}(k)$ to denote the cross- or auto-spectra monopole in real space and keep the subscript s to denote statistics computed in redshift space, together with the superscript ℓ , which refers to the multipole. To characterise spectra in redshift space, we rely on the three first non-zero multipoles, namely $P_{s,\alpha\beta}^{\ell=0}(k)$, $P_{s,\alpha\beta}^{\ell=2}(k)$, and $P_{s,\alpha\beta}^{\ell=4}(k)$, respectively, which are called the monopole, quadrupole, and hexadecapole. These $\ell \leq 4$ orders are the only non-vanishing moments in the linear approximation of the distortions and encode the full 2D information at linear scales (Kaiser 1987). The corresponding Legendre polynomials are

$$\mathcal{L}_\ell(\mu) = \begin{cases} 1 & \text{if } \ell = 0, \\ (3\mu^2 - 1)/2 & \text{if } \ell = 2, \\ (35\mu^4 - 30\mu^2 + 3)/8 & \text{if } \ell = 4. \end{cases} \quad (3)$$

Similarly to Bonnaire et al. (2021), we deconvolve the fields δ_α before estimating the power spectra to remove the biasing effect in the estimation introduced by the smoothing of the fourth-order B -spline interpolation to obtain the density fields. For additional robustness to any bias provoked by aliasing effects, the maximum Fourier bin, which we set at $k_{\max} = 0.5 h \text{ Mpc}^{-1}$ by default, is below half of the Nyquist frequency defined as $k_{\text{Nyq}} = \pi N_g/L = 0.57 h \text{ Mpc}^{-1}$.

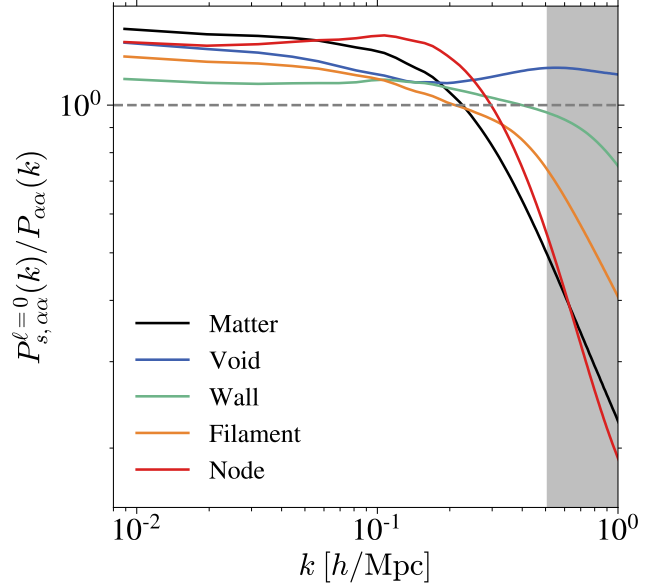


Fig. 1. Ratio between the redshift-space and real-space monopoles computed in the cosmic web environments. The grey shaded region shows the scales $k > 0.5 h \text{ Mpc}^{-1}$, which are unused in this analysis.

3.3. Cosmic web environments and redshift space

Figure 1 shows the ratios between the real-space and redshift-space monopoles obtained from the average of 7000 fiducial simulations. It is worth mentioning that the spectra are not weighted by the mass fraction of the environments, which in particular means that the matter multipoles cannot be recovered by summing the auto- and cross-spectra obtained in the environments, that is, $P_{s,\text{mm}}^\ell(k) \neq \sum_{\alpha\beta} P_{s,\alpha\beta}^\ell(k)$. As expected, all the spectra are boosted at large scales (small k values) due to the coherent motion of matter escaping from voids and moving towards dense regions because of the Kaiser effect. On the other hand, a decrease in power is observed at smaller scales where the FoG effect dominates, spreading particles initially residing in spherical overdensities along the line of sight. FoG mostly impacts the node environment where the highest velocities are found. Beyond the shape of the spectra in redshift space, the individual impact of each cosmological parameter is also different, as illustrated in Fig. 2. As an example, decreasing σ_8 does not imply a simple shift of the matter power spectrum monopole in redshift space but is boosting the power at small scales (large k values) due to the FoG effect. This is illustrated by the dashed lines of the top-left panel, which show the monopole residuals $P_{s,\text{mm}}^{\ell=0}(k)^{\theta_i} / P_{s,\text{mm}}^{\ell=0}(k)^{\text{fid}} - 1$ with θ_i being either σ_8^- , Ω_m^- , or M_ν^+ . These effects can then cause degeneracies between cosmological parameters. It has been shown for instance that the effect of massive neutrinos can be mimicked by a decrease in σ_8 on the redshift-space monopole at small scales $k > 0.1 h \text{ Mpc}^{-1}$ (see e.g. Villaescusa-Navarro et al. 2018; Hahn et al. 2020; Bayer et al. 2022), which in turn has a similar impact to a decrease in Ω_m . Comparing with the bottom-left panels of Fig. 2, we see that the monopoles computed in the cosmic web environments have various dependencies when decreasing the three parameters. Consequently, a change in σ_8 can no longer be reproduced by increasing the summed neutrino mass, illustrating the breaking of degeneracies that we expect by analysing the several environments separately. However, the anisotropic nature of the matter density field in redshift space leads to nonzero $\ell > 0$ multipoles that already have

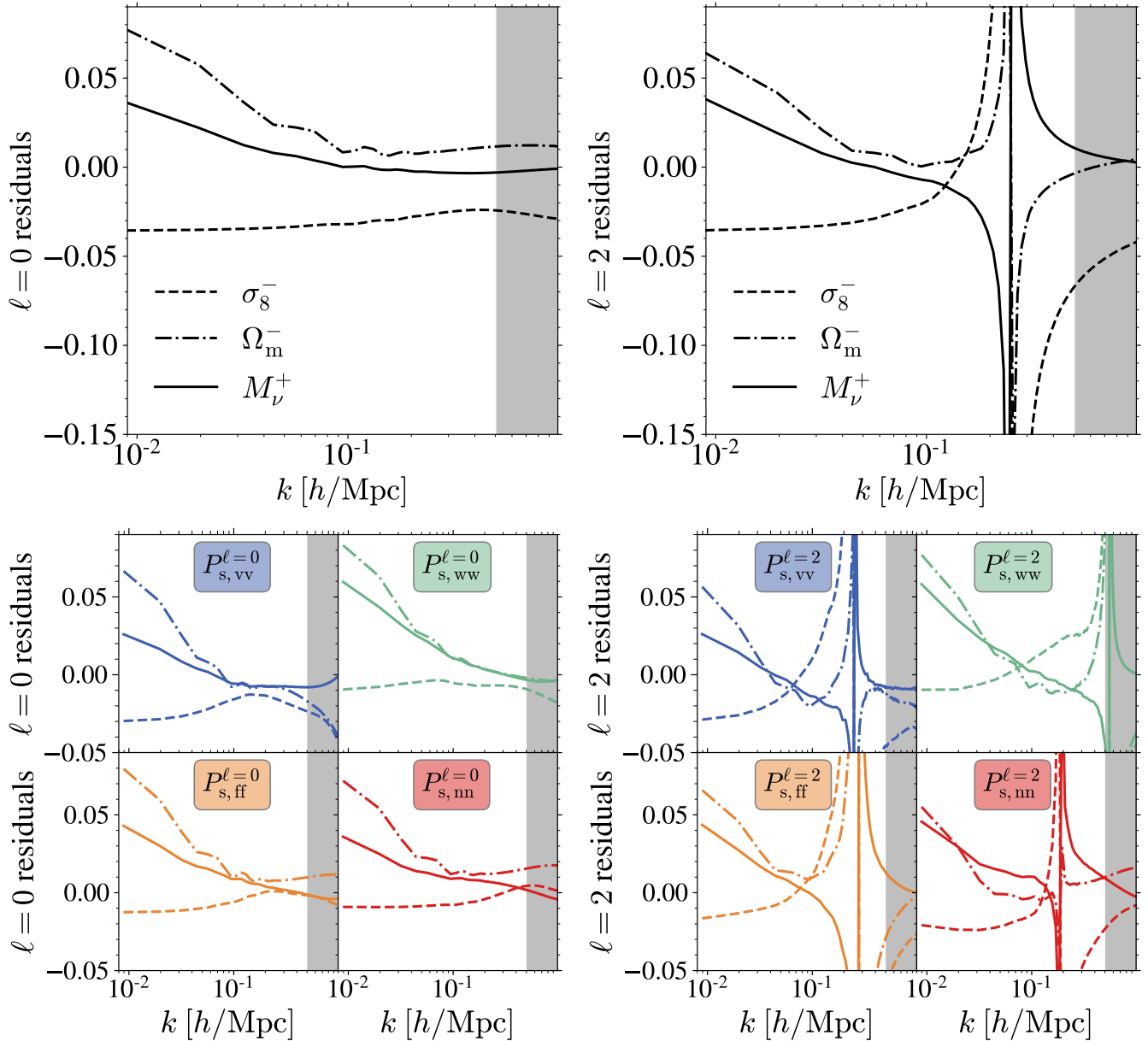


Fig. 2. Residuals of the matter monopole (upper left panel) or quadrupole (upper right panel) and environment-dependent monopoles (lower left panel) or quadrupoles (lower right panel) when varying either σ_8 , Ω_m , or M_ν . Residuals are defined as $P_{s,\alpha\alpha}^\ell(k)^{\theta_i}/P_{s,\alpha\alpha}^\ell(k)^{\text{fid}} - 1$ with $\theta_i = M_\nu^+$ (solid line), σ_8^- (dashed line), or Ω_m^- (dashed dotted line).

various shapes when varying these parameters, as illustrated in the upper right panel of Fig. 2 for $\ell = 2$. This suggests that combining the two first multipoles of the matter density field should allow us to break some degeneracies already. However, the measurements of the multipoles are noisier when going to larger ℓ and it might be interesting to measure the achievable accuracy of the environment monopoles only. The right-bottom panels also show that the environments themselves have different shapes for $\ell = 2$, suggesting further potential for imposing tighter constraints using the statistics drawn from combining the four monopoles and the four quadrupoles derived in the different environments.

The effects of the distortions are also clearly rooted in the classification itself, where we expect for instance the FoG to lead to a leakage of some cells classified in nodes in real space to filaments in redshift space. These effects are quantified in Table 1, which reports the confusion matrix averaged over ten simulations at fiducial cosmology when considering the real-

Table 1. Averaged confusion matrix obtained from ten boxes of the Quijote simulations with the fiducial cosmology when considering the absence of RSD as a ground truth.

Real space	Redshift space			
	Void	Wall	Filament	Node
Void	0.813	0.179	0.007	0.001
Wall	0.018	0.809	0.166	0.006
Filament	0.001	0.082	0.809	0.107
Node	0.002	0.042	0.367	0.587

space classification as the ground truth. The non-diagonal terms of this matrix therefore correspond to the fraction of change in

Table 2. Marginalised 1σ constraints obtained from the analysis of power-spectra monopoles and quadrupoles computed in the different environments for all cosmological parameters.

Statistics	σ_{Ω_m}	σ_{Ω_b}	σ_h	σ_{n_s}	σ_{σ_8}	σ_{M_v}
P_{mm}	0.0969	0.0413	0.5145	0.5019	0.0132	0.8749
$P_{\text{s,mm}}^{\ell=0}$	0.0081	0.0144	0.1495	0.0786	0.0115	0.3611
$P_{\text{s,mm}}^{\ell=\{0,2\}}$	0.0046	0.0133	0.1396	0.0719	0.0020	0.0834
$P_{\text{s,comb}}^{\ell=0}$	0.0033 (1.4)	0.0105 (1.3)	0.0877 (1.6)	0.0346 (2.1)	0.0026 (0.8)	0.0482 (1.7)
$P_{\text{s,comb}}^{\ell=\{0,2\}}$	0.0027 (1.7)	0.0097 (1.4)	0.0773 (1.8)	0.0295 (2.4)	0.0020 (1)	0.0304 (2.7)
$P_{\text{s,comb}}^{\ell=\{0,2\}} + P_{\text{s,mm}}^{\ell=\{0,2\}}$	0.0011 (4.1)	0.0091 (1.5)	0.0716 (2.0)	0.0279 (2.6)	0.0015 (1.4)	0.0151 (5.5)

Notes. Improvement factors (in parentheses) are relative to the matter case in redshift space, namely $P_{\text{s,mm}}^{\ell=\{0,2\}}$ in the third row. σ_{M_v} is in units of eV.

the environment classification when going from real to redshift space and a perfect match would lead to the identity matrix. The matrix shows for instance that 36.7% of the particles classified in nodes in real space are found in filaments when the classification is performed in redshift space. We note that the leakage mostly occurs between one environment and another of directly higher local dimension (i.e. voids to walls, walls to filaments, and filaments to nodes), while very little leakage occurs between environments of higher dimensions. We interpret this effect as a result of our plane-parallel approximation, which forces the distortions to occur along one spatial axis only. It is worth underlining that, apart from nodes, the classification is quite stable to the distortions, with about 80% of the particles in each environment (voids, walls, and filaments) keeping the same classification in real and redshift spaces.

4. Cosmological constraints

To quantify the amount of information carried by the spectra computed in the cosmic web environments, we make use of a Fisher analysis (details can be found in Appendix A). Briefly, this latter allows us to derive a lower bound on the constraints that one would achieve when relying on a given statistic, in our case either the spectra in the cosmic web environments or the full matter power spectrum. The derivation of the constraints mostly relies on two elements, both of which we derive numerically from the Quijote simulations: (i) the covariance matrix between Fourier amplitudes and (ii) the derivatives of the statistic with respect to the six cosmological parameters Ω_m , Ω_b , h , n_s , σ_8 , and M_v . The covariance matrices used in the Fisher forecast are shown in Appendix B.

The first row of Table 2 gives the constraints obtained from the real-space matter power spectrum presented in Bonnaire et al. (2021), while the second and third rows show the redshift-space constraints from the matter monopole $P_{\text{s,mm}}^{\ell=0}$ and the combination of the monopole and quadrupole $P_{\text{s,mm}}^{\ell=\{0,2\}}$, respectively. It is particularly striking that the redshift-space constraints obtained from just the monopole allow us to improve those in real space for Ω_m by a factor of up to 12 and those for M_v by a factor of up to 2.4. These gains even reach 21.1 and 10.5, respectively, when adding the information of the quadrupole in the $P_{\text{s,mm}}^{\ell=\{0,2\}}$ statistic. Such results are expected when moving from real to redshift space due to the additional velocity information coming into play in the breaking of the density field isotropy. It is also interesting to note that only σ_{σ_8} and σ_{M_v} are drasti-

cally reduced when adding the quadrupole information to the matter monopole, while the error on other parameters remains roughly the same, quantifying the expected breaking of degeneracy between these two parameters discussed in Sect. 3.3 and illustrated in the top panels of Fig. 2.

Concerning the environment-dependent spectra, we report the following in rows four to six of Table 2:

- Using the combination of environment monopoles (fourth line of Table 2) already provides tighter constraints than the combination of the first two multipoles of the full matter density field, $P_{\text{s,mm}}^{\ell=\{0,2\}}$ (third line of Table 2), with improvement factors of up to 2.1 and 1.7 for n_s and M_v , respectively.

- Adding the $\ell = 2$ information brings down the constraints to even lower values with improvement factors of $\{1.7, 1.4, 1.8, 2.4, 1, 2.7\}$ on cosmological parameters $\{\Omega_m, \Omega_b, h, n_s, \sigma_8, M_v\}$ with respect to the matter analogue in redshift space, as shown in the fifth row of Table 2. Similarly to what is observed in the matter case, adding up the $\ell = 2$ information mostly breaks the degeneracy on M_v without greatly impacting the set of other parameters.

- The constraints imposed by the combination of environment statistics highly depend on the maximum scale of the spectra included in the analysis. Many panels of Fig. 3 indeed indicate that the matter monopole+quadrupole saturates at mildly non-linear scales, a feature already exhibited in the real space analysis from Bonnaire et al. (2021) and previously pointed out by Takahashi et al. (2010), Blot et al. (2015), Chan & Blot (2017). This saturation, which is mainly due to the degeneracies among parameters of the cosmological model, is not observed for the combination of environment-dependent spectra, and is a particular encouragement for us to push the limiting scale of the analyses further in order to fully capitalise on the potential of the cosmic web split.

- Unlike the real-space case, the improvement factors of the individual environments are below one for all parameters when $k_{\text{max}} = 0.5 h \text{Mpc}^{-1}$ (see Table C.1). This observation is not true at all scales, as illustrated by Fig. 3, where the environments perform individually better than the matter statistics when restricting the analysis to very large scales, that is, below $0.15 h \text{Mpc}^{-1}$ in practice, except for σ_8 and M_v for which the full matter is always more attractive than individual environments.

- The absolute values of the constraints derived from individual environments and their combination in redshift space are similar to the constraints derived from real space (comparing Table C.1 to Table C.1 in Bonnaire et al. 2021), and are even tighter for Ω_m and M_v .

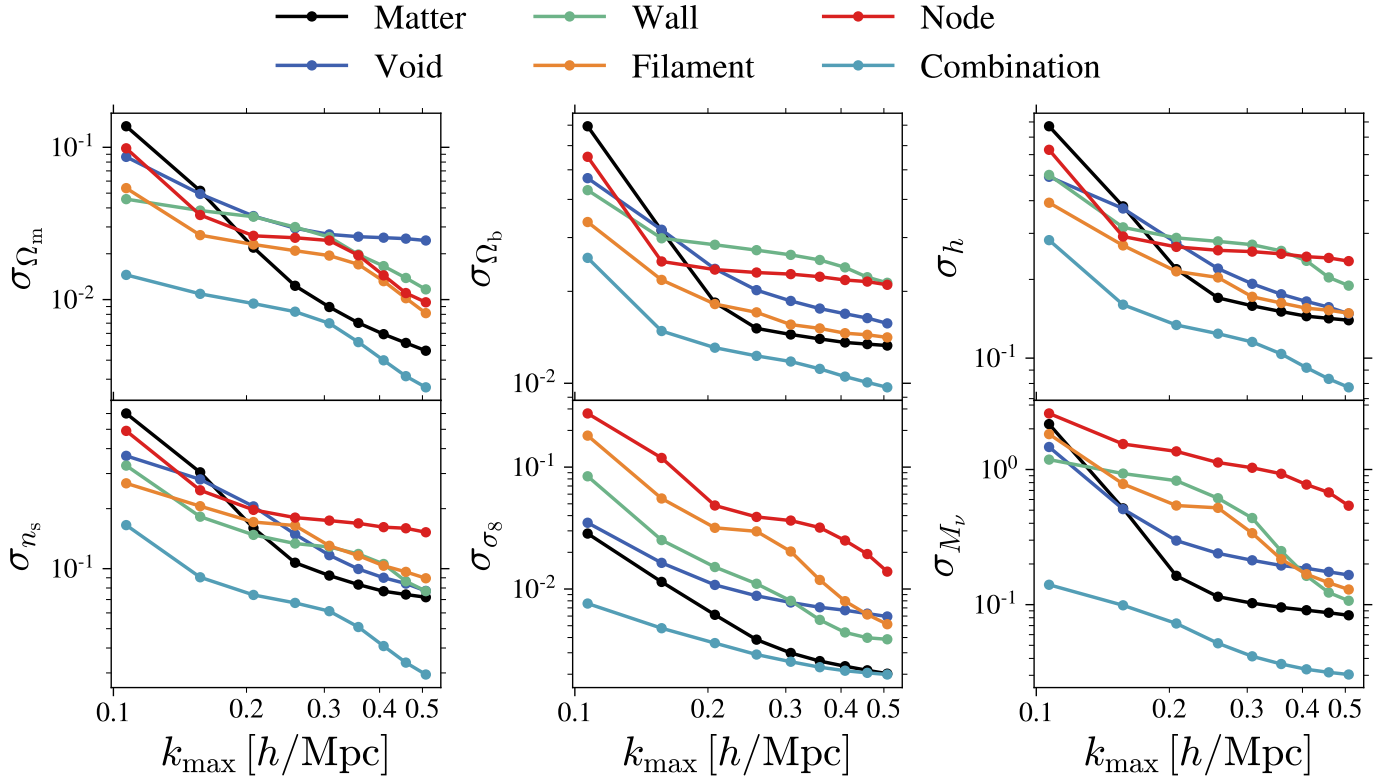


Fig. 3. Evolution of the marginalised constraint σ_{θ_i} given by $P_{\alpha\alpha}^{\ell=\{0,2\}}$ on cosmological parameters $\{\Omega_m, \Omega_b, h, n_s, \sigma_8\}$ and the sum between neutrino mass M_ν and the maximum scale used for the Fisher analysis, k_{\max} . Here, σ_{M_ν} is in units of eV.

– The non-negligible improvements of the constraints drawn from $P_{s,\text{comb}}^{\ell=\{0,2\}}$ are amplified when the analysis is reduced to large scales only (i.e. small values of k), especially for parameters such as Ω_m , Ω_b , σ_8 , and M_ν , reaching a factor 5 improvement over the matter statistic, as represented in Fig. 3.

– The overall information brought by the combination of environment-dependant spectra can still be completed by the matter multipoles to further improve the constraint by factors of up to 5.5 and 4.1 on M_ν and Ω_m , respectively (see the last row of Table 2). From inspection of the confidence ellipses in Fig. B.2, it indeed appears that $P_{s,\text{comb}}^{\ell=\{0,2\}}$ and $P_{s,\text{mm}}^{\ell=\{0,2\}}$ bring complementary information in relations such as $M_\nu - \sigma_8$ or $\sigma_8 - \Omega_m$, which explains the tighter constraints obtained when merging this information in the analysis.

– Including the cross spectra in the combination, and therefore considering the statistic $P_{s,\alpha\beta}^{\ell=\{0,2\}}$ for $(\alpha, \beta) \in \{v, w, f, n\}^2$, allows us to boost the gains on some parameters, such as M_ν and Ω_m by factors of 2.2 and 3.5, respectively. All these gains are summarised in Fig. 4 and this statistic notably involves the information from auto-spectra and accounts for a total of 800 elements in the summary statistic.

We also checked that including the hexadecapole ($\ell = 4$) in our analysis does not lead to any further improvement of the constraints over the $\ell = \{0, 2\}$ statistics, neither for the matter nor for the environment-dependant spectra, or their combination.

5. Discussion and conclusion

It is well known that by analysing the matter distribution in redshift space through its power spectrum, we can greatly improve cosmological constraints. This is thanks to the velocity information contained in the anisotropy of the density fields, which

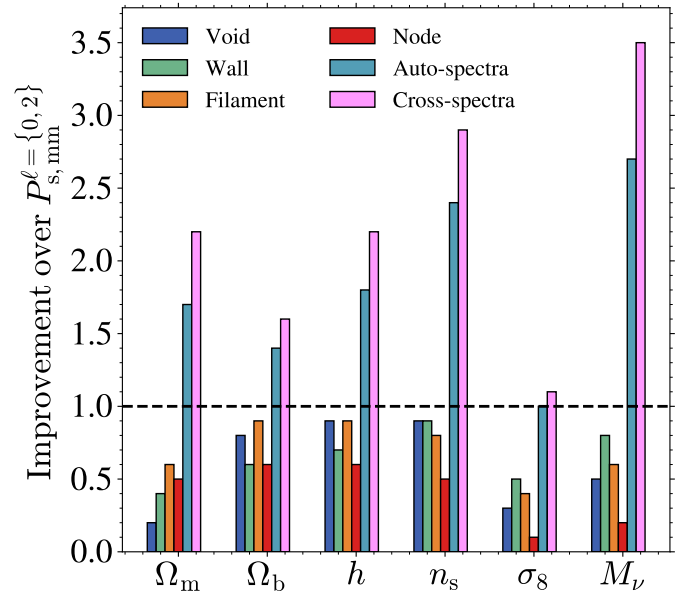


Fig. 4. Improvement factors of the several studied statistics over the redshift-space matter monopole+quadrupole constraints for each of the six cosmological parameters at $k_{\max} = 0.5 h\text{Mpc}^{-1}$. The horizontal black line shows the unity improvement. We note that these statistics exclude the combination with the matter multipoles and only concern the several cosmic web environments and their combination (therefore excluding the last line of Table 2 for instance).

breaks key degeneracies between parameters. Our results show that we can attain even tighter constraints by performing statistical analyses in identified cosmic web environments. In this

configuration, the constraints on cosmological parameters, such as the initial spectral index, the matter density, the Hubble parameter, and most notably the summed neutrino mass, can be improved by up to a factor of two. Improvement factors can even reach five when combining these analyses with full matter statistics. Our approach therefore opens up a new avenue for the derivation of cosmological constraints via the analysis of the observed distribution of tracers. In particular, it shows that conditioning the correlations on the cosmic web environments improves constraints on the cosmological parameters compared to the direct correlation of all tracers, in agreement with previous findings (e.g. Uhlemann et al. 2020; Paillas et al. 2021, 2023).

The present study was performed in a theoretical context in which we built matter density fields from numerical simulations. Firstly, within this idealised context and with the studied statistics, it appears that constraining the cosmological parameters in redshift space yields similar constraints to those obtained in real space, if not tighter, despite the fact that we do not handle the distortions in the classification. This therefore suggests that reconstruction of the particles from redshift to real space (e.g. Jasche & Wandelt 2013; Bos et al. 2014; Leclercq et al. 2015) is not necessary in the present setting and can even deteriorate the accuracy of the derived constraints. Secondly, this context allows us to explore the full range of scales including non-linear ones. The combination of auto-spectra computed in the different environments up to $k_{\max} = 0.5 h \text{Mpc}^{-1}$ leads to tighter constraints in redshift than real space for most parameters, which is even more emphasised when adding information from the full matter monopole and quadrupole.

With this second paper of the series, we therefore go beyond showing how to use the cosmic web environments as leverage in order to improve the cosmological constraints in real space, and present the benefits of using a cosmic web classification to improve redshift-space constraints over the traditional matter power-spectrum multipoles. In a third paper, we will complete our comprehensive analysis of the cosmological content of the cosmic web environments with the analysis of the direct higher order statistic, namely the bispectrum. The next steps and challenges will consist in adapting and optimising our analysis – which is currently performed in a theoretical and idealised setup of large number density provided by particles in simulations – to more realistic cases of observed sparse tracers of the density field, such as halos or galaxies, as performed in other works (such as Hahn et al. 2020; Naidoo et al. 2022; Kreisch et al. 2022; Paillas et al. 2023). In this context, simulations will represent a major endeavour. Adapting our analysis so that it can be applied to actual data from galaxy surveys requires the exploitation of simulations with a higher mass resolution, enabling the definition of galaxy-type halos and with realistic baryonic physics models Vogelsberger et al. (2014). Such simulations will also allow us to bypass the difficulty of a theoretical modelling of covariance matrices and the building of likelihoods, which can be achieved with the use of simulation-based inference (e.g., Cranmer et al. 2020; Lemos et al. 2023). The information that can be gained on the cosmological content of cosmic web environments, as demonstrated by the results of the present study, will certainly justify the effort involved. The work presented here therefore constitutes a step towards an optimal and interpretable summary statistics that will allow us to take full advantage of the future Stage IV experiments (Laureijs et al. 2011; Spergel et al. 2015; DESI Collaboration 2016).

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Appendix A: Fisher formalism

The Fisher formalism allows the forecast of constraints we will get on our set of parameters $\theta = \{\Omega_m, \Omega_b, h, n_s, \sigma_8, M_\nu\}$ based on a statistical summary of the density fields. In our case, this summary is given by the set of environment-dependent (cross- or auto-)power spectra in Fourier bins, $\mathbf{s} = \{P_{\alpha\beta}(k)\}$. Assuming a Gaussian likelihood with a covariance matrix that is independent of cosmology (Carron 2013; Kodwani et al. 2019), we can derive elements of the Fisher information matrix as

$$[\mathbf{I}(\theta)]_{i,j} = \left(\frac{\partial \bar{\mathbf{s}}}{\partial \theta_i} \right)^\top \boldsymbol{\Sigma}^{-1} \left(\frac{\partial \bar{\mathbf{s}}}{\partial \theta_j} \right). \quad (\text{A.1})$$

The first ingredient in the computation of the Fisher information matrix is the unbiased estimate of the inverse covariance matrix $\boldsymbol{\Sigma}^{-1}$, which is given, still under the Gaussian assumption, by (Kaufman 1967; Hartlap et al. 2007)

$$\boldsymbol{\Sigma}^{-1} = \frac{N_{\text{fid}} - n - 2}{N_{\text{fid}} - 1} \hat{\boldsymbol{\Sigma}}^{-1}, \quad (\text{A.2})$$

where N_{fid} is the number simulations at the fiducial cosmology, n is the length of the summary statistics vector \mathbf{s} , and $\hat{\boldsymbol{\Sigma}} = (\mathbf{s} - \bar{\mathbf{s}})(\mathbf{s} - \bar{\mathbf{s}})^\top / (N_{\text{fid}} - 1)$ is the unbiased estimate of the covariance matrix.

The second ingredient of Eq. (A.1) are the derivatives with respect to the parameters of the model θ . These latter can be estimated numerically from the $N_{\text{deriv}} = 500$ realisations of the variation of individual parameters in the Quijote suite as

$$\frac{\partial \bar{\mathbf{s}}}{\partial \theta_i} \simeq \frac{\bar{\mathbf{s}}(\theta_i + d\theta_i) - \bar{\mathbf{s}}(\theta_i - d\theta_i)}{2d\theta_i}. \quad (\text{A.3})$$

However, we note that the above definition does not apply for M_ν being a positive quantity with a fiducial value at 0.0 eV. For this parameter, we therefore rely on the four-point forward approximation of the derivative

$$\frac{\partial \bar{\mathbf{s}}}{\partial M_\nu} \simeq \frac{\bar{\mathbf{s}}(4M_\nu^+) - 12\bar{\mathbf{s}}(2M_\nu^+) + 32\bar{\mathbf{s}}(M_\nu^+) - 21\bar{\mathbf{s}}(M_\nu = 0.0)}{12M_\nu^+}, \quad (\text{A.4})$$

where $M_\nu^+ = 0.1$ eV. For consistency, we use a set of simulations initialised using the Zel'dovich approximation to compute $\bar{\mathbf{s}}(M_\nu = 0)$ in the previous equation because all the other terms are initialised this way. The presented results are obtained by making use of $N_{\text{deriv}} = 500$ and $N_{\text{fid}} = 7000$ realisations to respectively compute the derivatives and the covariance matrix.

Appendix B: Covariance matrices, derivatives, and confidence ellipses

While discussed in the main text Sect. 4, we report in this section the two ingredients of the Fisher quantification of information, which are: (i) the correlation matrices between the several elements of the statistic vector in Fig. B.1 and (ii) the derivatives of the redshift-space monopoles and quadrupoles with respect to the six cosmological parameters. Figure B.2 shows the corner plot with the several confidence ellipses obtained for the $P_{s,\alpha\alpha}^{\ell=(0,2)}$ statistics for all six cosmological parameters, namely Ω_m , Ω_b , h , n_s , σ_8 , and M_ν .

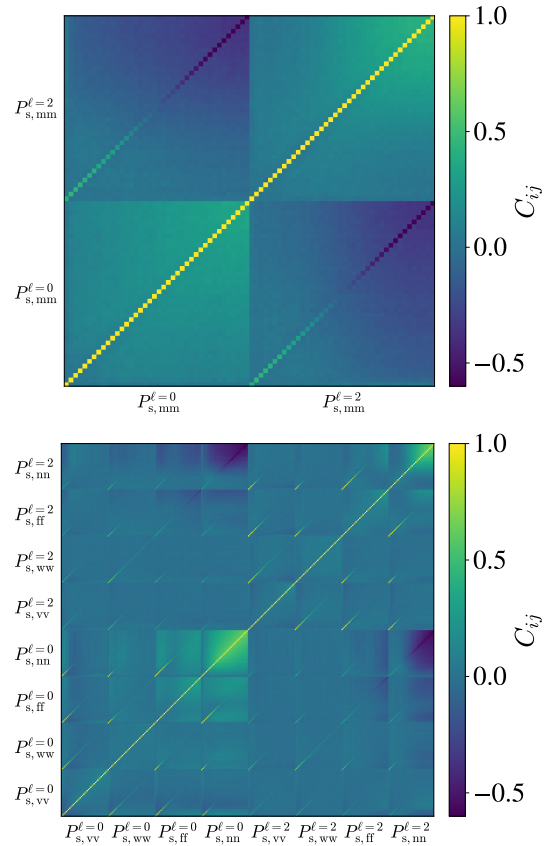


Fig. B.1. Correlation coefficients C_{ij} obtained from $N_{\text{fid}} = 7000$ simulations for (top) the matter power spectrum, and (bottom) from the auto-spectra $P_{\alpha\alpha}(k)$ computed in the different environments. Each submatrix goes from $k = 0.1$ h/Mpc to $k = 0.5$ h/Mpc.

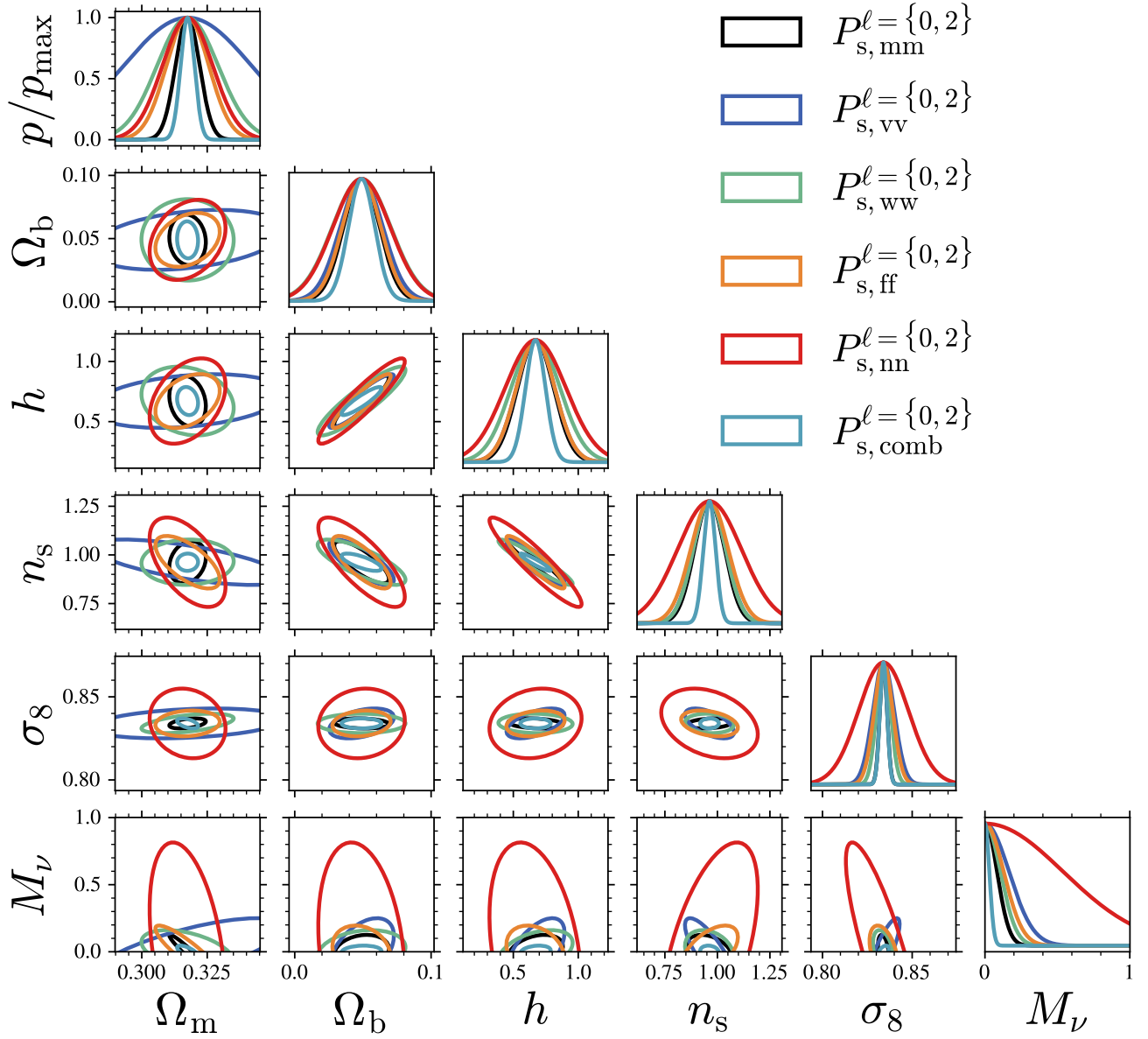


Fig. B.2. 1σ confidence ellipses for all the pairs of cosmological parameters ($\Omega_m, \Omega_b, h, n_s, \sigma_8, M_\nu$) obtained from the several statistics (either the matter monopole+quadrupole or the ones from the environments and their combination) in redshift space. The normalised probability density functions for each parameter are shown on the diagonal.

Appendix C: Constraints from spectra in the several environments

Here, we provide the individual constraints obtained from the monopole and quadrupole power spectra computed in the different cosmic web environments in Table C.1. As discussed

in Sect. 4, they all lead to much weaker constraints than the full matter statistic in redshift-space when $k_{\max} = 0.5 h/\text{Mpc}$. However, their combination leads to tighter constraints with an improvement by a factor of up to 2.7 for the summed neutrino mass and 2.4 for n_s .

Table C.1. Marginalised 1σ constraints obtained from the analysis of power-spectra monopoles and quadrupoles computed in the different environments for all cosmological parameters.

Statistics	σ_{Ω_m}	σ_{Ω_b}	σ_h	σ_{n_s}	σ_{σ_8}	σ_{M_ν}
$P_{s,vv}^{\ell=\{0,2\}}$	0.0244 (0.2)	0.0157 (0.8)	0.1478 (0.9)	0.0773 (0.9)	0.0060 (0.3)	0.1657 (0.5)
$P_{s,ww}^{\ell=\{0,2\}}$	0.0117 (0.4)	0.0213 (0.6)	0.1893 (0.7)	0.0772 (0.9)	0.0039 (0.5)	0.1069 (0.8)
$P_{s,ff}^{\ell=\{0,2\}}$	0.0081 (0.6)	0.0141 (0.9)	0.1485 (0.9)	0.0895 (0.8)	0.0051 (0.4)	0.1294 (0.6)
$P_{s,nn}^{\ell=\{0,2\}}$	0.0096 (0.5)	0.0210 (0.6)	0.2350 (0.6)	0.1523 (0.5)	0.0139 (0.1)	0.5397 (0.2)
$P_{s,comb}^{\ell=\{0,2\}}$	0.0027 (1.7)	0.0097 (1.4)	0.0773 (1.8)	0.0295 (2.4)	0.0020 (1)	0.0304 (2.7)

Notes. All the improvement factors (in parentheses) are relative to the matter case in redshift-space, namely $P_{s,mm}^{\ell=\{0,2\}}$ given in the third row of Table 2. σ_{M_ν} is in units of eV.