



HAL
open science

Olivine of carbonaceous chondrites and reservoir of nebular hydrogen

Samantha Azevedo-vannson, Laurent Remusat, Laurette Piani, F. Danoix, Hugues Leroux, Sylvain Bernard, Jean-christophe Viennet, Sylvain Pont, Raphaële Danoix, Maximilien Verdier-Paoletti, et al.

► **To cite this version:**

Samantha Azevedo-vannson, Laurent Remusat, Laurette Piani, F. Danoix, Hugues Leroux, et al.. Olivine of carbonaceous chondrites and reservoir of nebular hydrogen. 85th Annual Meetings of the Meteoritical Society, Meteoritical Society, Aug 2022, Glasgow, United Kingdom. hal-03856226

HAL Id: hal-03856226

<https://hal.science/hal-03856226v1>

Submitted on 9 Feb 2024

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

OLIVINE OF CARBONACEOUS CHONDRITES AND RESERVOIR OF NEBULAR HYDROGEN.

Samantha Azevedo-Vannson¹, Laurent Remusat¹, Laurette Piani², Frédéric Danoix³, Hugues Leroux⁴, Sylvain Bernard¹, Jean-Christophe Viennet¹, Sylvain Pont¹, Raphaele Danoix³, Maximilien Verdier-Paoletti¹, Roger Hewins¹, Mathieu Roskosz¹, ¹ Muséum National d'Histoire Naturelle, Institut de Minéralogie, de Physique des Matériaux, et de Cosmochimie (IMPMC), UMR CNRS 7590, ²Université de Lorraine, CNRS, CRPG, F-54000 Nancy, France, ³Groupe de Physique des Matériaux, UMR CNRS 6634, Faculté des Sciences de Rouen, Avenue de l'Université, F-76801 Saint Etienne du Rouvray, France, ⁴Univ. Lille, CNRS, INRAE, ENSCL, UMR 8207 - UMET - Unité Matériaux et Transformations, F-59000 Lille, France

Introduction: Olivine is a nominally anhydrous mineral usually depleted in water unless formed in a high pressure and temperature environment^[1]. Several studies have found 10 to 100 times more water in olivine of carbonaceous and ordinary chondrites than on Earth^[2,3]. In the meantime, Shimizu et al.^[4] detected less than 15 ppm of hydrogen in silicate minerals in chondrules of Semarkona (LL3), QUE 97008 (L3.05) and DOM 08006 (CO3).

Two main reservoirs of hydrogen, with distinct isotopic compositions, could likely contribute to chondrule mineral hydrogen content: the molecular gas of the protosolar nebula with a low D/H^[5] and water ice. The D/H ratio of water in solar system objects exhibit variations related to the distance to the star. Part of the water can thus have a presolar origin (high D/H ratio) and another part can result from re-equilibration with the H₂ of the disk^[6]. Water ice accreted on the parent body of hydrated meteorites like carbonaceous chondrites. Due to asteroid internal heating, ice melted and water percolated leading to aqueous alteration of its components. Both matrices and chondrules were affected. As a consequence, aqueous alteration could have contributed to the water budget of chondrules' olivine as well, therefore making the deconvolution of pre-accretionary and post-accretionary contributions difficult.

Samples and Methods: Three carbonaceous chondrites showing different alteration degrees were investigated: Paris (CM2.7 to CM2.9), Aguas Zarcas (CM2) and Mukundpura (CM2). Three polished sections were prepared in epoxy and carbon coated. Hydrogen content and isotopic composition of olivine in type I (FeO-poor) and type II (Feo-rich) chondrules and isolated olivine were analyzed by Cameca SIMS IMS 1280 HR2 (CRPG, Nancy) and by Cameca NanoSIMS 50 (MNHN, Paris) using an O- primary beam from a Hyperion radio-frequency plasma source in both cases. Olivine crystals isolated in the matrix are thought to be fragments of chondrules ejected during their formation. Thus, like the chondrule olivine, they have crystallized in the protoplanetary disc a few million years after the Solar System formation.

4 tips were prepared by focused ion beam (FIB) milling using a Thermofisher Helios G4 UXe DualBeam Xenon plasma FIB, and analysed by atom probe tomography (APT) using a Cameca LEAP 4000XHR instrument. The hydrogen contents (concentration, spaciation) were compared with those of anhydrous San Carlos olivine prepared and analysed in the same conditions with the same instruments.

Results: Olivine of carbonaceous chondrites contain large and variable amounts of non-labile hydrogen as attested by SIMS and atom probe tomography. The isotopic composition of this hydrogen varies between the signature of H₂ in the disk (D/H ~ 26 ± 7 × 10⁻⁶^[5]) and water ices (D/H > 156 × 10⁻⁶^[7]). The D/H ratios of Paris olivine cover a wide range of values (99 to 151 × 10⁻⁶) and are not correlated to Mg#. Mg-rich olivine of Aguas Zarcas also show this large variability (77 to 151 × 10⁻⁶). However, the D/H ratios is more homogeneous around 115 ± 10 × 10⁻⁶ when the iron content is above 40%. The olivine of Mukunpura, the most altered chondrite, have homogeneous isotopic compositions around 134 ± 7 × 10⁻⁶.

Discussion: With increasing degree of aqueous alteration on the parent body, the hydrogen content and D/H of olivine tends to be more homogeneous; D/H ratio of hydrogen in olivine tends to become isotopically akin to terrestrial mantle water^[11]. Despite the marks of some aqueous activity in these chondrites, the isotopic variabilities measured in Paris and Mg-rich Aguas Zarcas olivine might be an heritage of the different sources (ice or hydrogen of the protoplanetary disk) of hydrogen that embedded solid precursors during olivine formation.

Acknowledgments: This work was supported by the European Research Council via the ERC project HYDROMA (grant agreement No. 819587).

References: [1] Mackwell, S.J. and Kohlstedt, D.L. (1990) *Journal of Geophysical Research* 95:5079-5088. [2] Deloule, E. et al. (1998) *GCA* 62 :3367-3378. [3] Stephant, A. et al. (2017) *GCA* 199 :75-90. [4] Shimizu, K. et al. (2021) *GCA* 301 :230-258. [5] Mahaffy, P.R. et al. (1998) *Space Science Reviews* 84:251-263. [6] Jacquet, E. and Robert, F. (2013) *Icarus* 223 :722-732. [7] Hartogh, P. et al. (2011) *Nature* 478 :218-220. [11] Hallis, L.J. et al. (2015) *Science* 350 :795-797.