



HAL
open science

Narrow frequency-windowed autocorrelations as a diagnostic of solar-like stars

Ian W. Roxburgh

► **To cite this version:**

Ian W. Roxburgh. Narrow frequency-windowed autocorrelations as a diagnostic of solar-like stars. *Astronomy and Astrophysics - A&A*, 2009, 506, pp.435-444. 10.1051/0004-6361/200911906. hal-03785586

HAL Id: hal-03785586

<https://hal.science/hal-03785586>

Submitted on 6 Oct 2022

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Narrow frequency-windowed autocorrelations as a diagnostic of solar-like stars

I. W. Roxburgh^{1,2}

¹ Queen Mary University of London, Astronomy Unit, Mile End Road, London E1 4NS, UK
e-mail: I.W.Roxburgh@qmul.ac.uk

² LESIA, Observatoire de Paris, Place Jules Janssen, 92195 Meudon, France

Received 20 February 2009 / Accepted 17 June 2009

ABSTRACT

Aims. This paper investigates the diagnostic potential of narrow, frequency-windowed autocorrelation as a tool for probing the properties of solar-like oscillating stars when the determination of individual frequencies is impossible or is subject to large uncertainties, and when mode identification is difficult.

Methods. I use theoretical analysis including phase-shifts, modelling, and data analysis.

Results. Narrow-windowed autocorrelation of a time series can reveal the variation with frequency of the large separations $\Delta(\nu)$ and the half large separations Δ_{01}, Δ_{10} , thus helping with mode identification. This technique is applied to the CoRoT p-mode oscillators HD 49933, HD 175726, HD 181420, and HD 181906. Theoretical analysis and modelling are presented to illustrate the technique.

Key words. stars: oscillations – methods: analytical – methods: data analysis

1. Introduction

For solar-like stars reliable determining of p-mode frequencies from power spectra is not always possible since the amplitudes of the stochastically excited modes are very small, giving low signal/noise. For F stars observed by CoRoT¹ the line widths are large, which hinders the reliable determination of frequencies and for some stars, particularly HD 175726, individual frequencies are exceedingly difficult to extract. I here consider an alternative approach that may be useful when faced with poor quality data, namely the use of the autocorrelation of the time series. As shown by Roxburgh & Vorontsov (2006a) by adding noise to a solar power spectrum, this has diagnostic potential when faced with noisy data.

The autocorrelation of a discrete time series $G(t_k)$, $k = 0, N$ is

$$A(t_k) = \sum_{i=0}^N G(t_i) G(t_{i+k}). \quad (1)$$

One can see immediately that peaks are expected in the autocorrelation function of a photometric time series due to p-mode oscillations. Such oscillations are acoustic waves excited by surface convection. A wave packet produced near the surface of the star propagates to the far side of the star in time $2T$, where T is the acoustic radius of the star, and is reflected back arriving at (or near) the point of emission at $4T$. Thus one expects a peak in the autocorrelation at $4kT$, $k = 1, 2, 3 \dots$. Since the large separations $\Delta_{n,\ell} = \nu_{n,\ell} - \nu_{n-1,\ell}$ between modes with frequencies $\nu_{n,\ell}, \nu_{n-1,\ell}$

are approximately $1/(2T)$ the peaks in the autocorrelation give some mean value of Δ (cf. Roxburgh & Vorontsov 2006a).

To concentrate on frequency components in the photometric time series of a star in the range of p-mode oscillations, one can filter the time series by transforming to the frequency domain, windowing the resulting Fourier transform to retain only contributions in a given frequency range, then transform back to the time domain and take the autocorrelation of the resulting time series.

From the Wiener-Khinchin theorem the autocorrelation $A(t)$ is equal to the real part of the Fourier transform \mathcal{F} of the power spectrum $P(\nu)$ of the time series, and the amplitude envelope $A_e(t)$ by the amplitude of the full Fourier transform:

$$A(t) = \text{Re}(\mathcal{F}\{P(\nu)\}), \quad A_e(t) = |\mathcal{F}\{P(\nu)\}|, \quad P(\nu) = |\mathcal{F}\{G(t)\}|^2 \quad (2)$$

(cf. Press et al. 1992). It is therefore not necessary to transform back to the time domain but simply to take the Fourier transform of the filtered power spectrum. Figure 1 shows the autocorrelation and the amplitude envelope of the photometric time series of the F5V star HD 49933 observed for 60 days in the initial run of CoRoT (see below and Appourchaux et al. 2008), which has been filtered (using a \sin^2 filter) in the frequency interval $\nu = 1200\text{--}2400 \mu\text{Hz}$. The amplitude envelope $A_e(t)$ removes the rapidly oscillating component in the autocorrelation due to the frequencies.

To determine the location of the peaks in the filtered autocorrelation function (Fig. 1), one could fit a Gabor function to the rapidly oscillating autocorrelation (cf. Kholikov & Hill 2008), but it is more convenient to use the amplitude envelope or its square, the autocorrelation power. As mentioned above, the first peak in the autocorrelation is expected at $4T$, where $T = \int dr/c$ is the acoustic radius of the star, the time for a wave to travel from centre to surface (here $c(r)$ is the sound speed), and this

¹ The CoRoT space mission, launched on 2006 December 27, was developed and is operated by CNES, with the participation of the Science Programmes of ESA, ESA's RSSD, Austria, Belgium, Brazil, Germany and Spain.

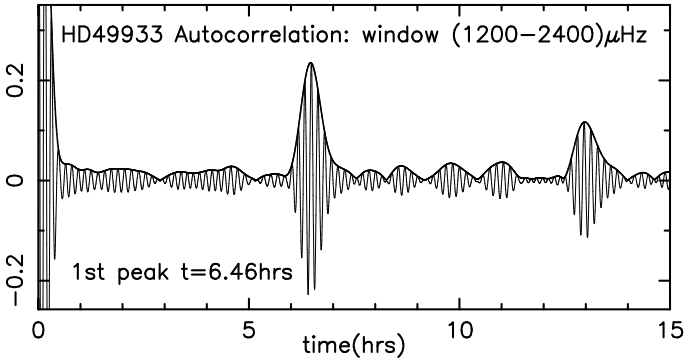


Fig. 1. HD 49933 autocorrelation, frequency windowed between 1200–2400 μHz , and its amplitude envelope.

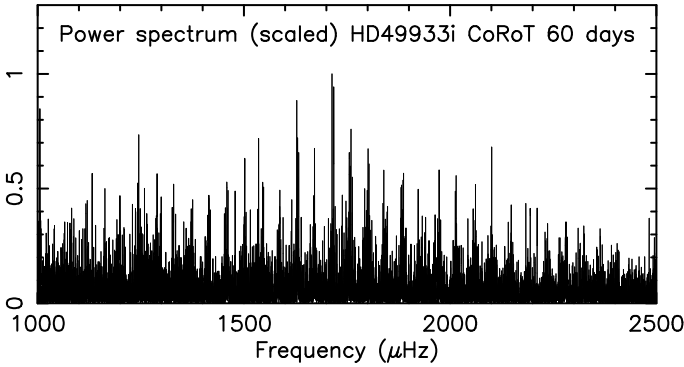


Fig. 2. HD 49933 power spectrum.

gives an estimate of the large separation Δ . However, as the actual ray path and the location of the reflecting layer in the atmosphere depend on the star's structure and on frequency and degree, the round trip travel time is not independent of ν, ℓ , and the large separations $\Delta_{n,\ell}$ vary with both frequency and degree. This variation is primarily due to the quasi-periodic modulation of the frequencies caused by the HeII ionisation layer and to the different contributions of the inner layers to frequencies of different degree ℓ . Additionally there can be a small contribution at $2T$ due to reflection from the core (Roxburgh & Vorontsov 1996,7). I here examine whether information on this variation with frequency and degree can be extracted from the autocorrelation power by narrow windowing in the frequency domain.

2. The CoRoT star HD 49933

The F5V ($m_v = 5.77$) star HD 49933 was observed for 60 days in the initial run of CoRoT (see Appourchaux et al. 2008). After applying several corrections, details of which are given in Samadi et al. (2006), the light curves are sampled at a cadence of 32 s, the duty cycle is $\sim 90\%$ and the gaps, mainly due to the passages of the satellite through the South Atlantic anomaly, are filled by interpolation. (This is level 2 data in the language of CoRoT.) The typical length of the gaps is ~ 9 min, producing a $\sim 10\%$ reduction in power at frequencies > 2000 μHz (see Appourchaux et al. 2008).

The resulting light curve has a mean drift and significant variations due to rotation and activity, this can be removed by subtracting off a running mean. Linear interpolation in the gaps can be modified by adding noise based on the mean variation outside the gaps. Experiments using different running means, and none, and different gap-filling procedures showed that the

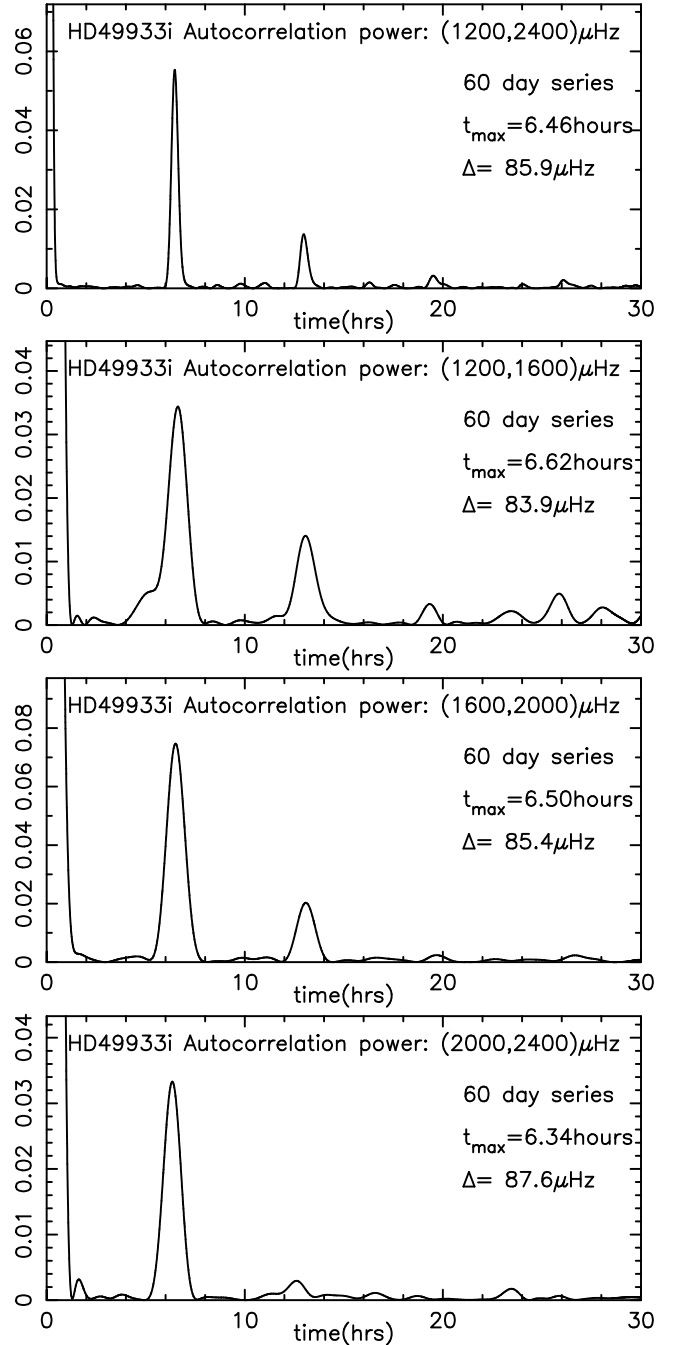


Fig. 3. Frequency windowed autocorrelation power for HD 49933 (scaled): **a)** windowed between 1200–2400 μHz , **b)** narrow window 1200–1600 μHz , **c)** 1600–2000, **d)** 2000–2400.

results described below were independent of these procedures. The power spectrum using the level 2 data is shown in Fig. 2 for the frequency range 1000–2500 μHz .

This power spectrum was then filtered with a \sin^2 window $W(\nu) = \sin^2(\pi(\nu - \nu_1)/(\nu_1 - \nu_2))$ with $(\nu_1, \nu_2) = (1200, 2400)\mu\text{Hz}$ and the resulting autocorrelation (normalised to 1 at $t = 0$) is that shown Fig. 1 above, and the autocorrelation power in Fig. 3a. The first peak occurs at a time $t_m = 6.46$ h corresponding to a mean large separation of 85.9 μHz , the second peak at $2t_m$ and a third peak can just be detected at $3t_m$. The rapid decrease in the height of successive peaks in Fig. 3a indicates the short life time of the modes; this is to be expected since the line widths in the power spectrum are broad (Appourchaux et al. 2008).

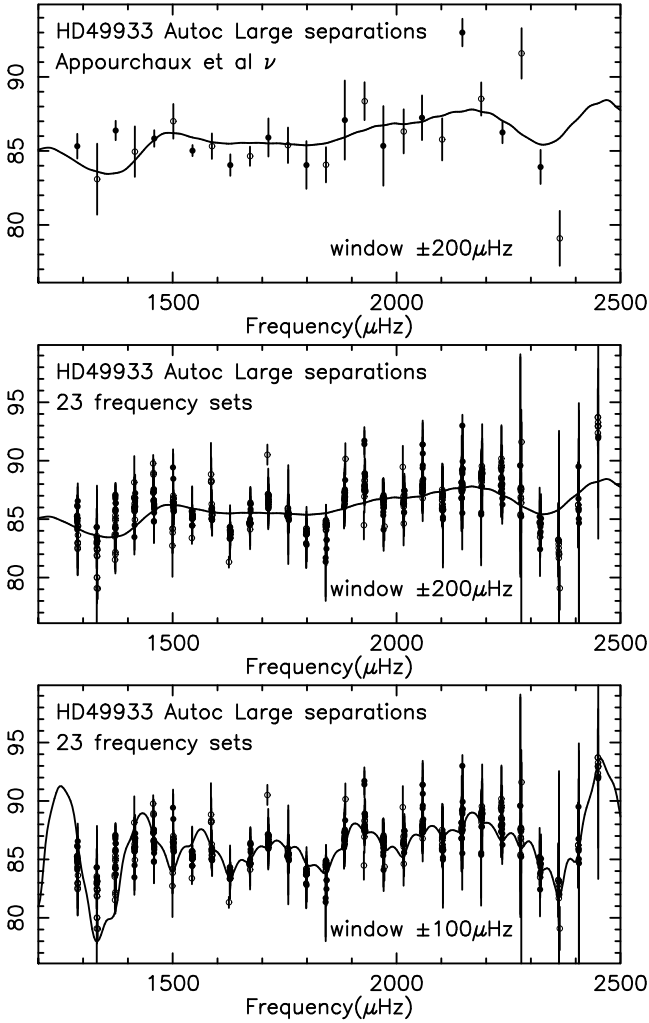


Fig. 4. a) Variation of large separation with frequency for HD 49933 from autocorrelation with a \sin^2 frequency window of full width $400 \mu\text{Hz}$, and the large separations with 1σ error bars as determined by Appourchaux et al. (2008). b) As in a) but with large separations from 23 different frequency determinations. c) Autocorrelation large separations with a narrow window of $200 \mu\text{Hz}$.

Figures 3b–d show the autocorrelation power for three independent narrower windows of $400 \mu\text{Hz}$. These again show the characteristic maxima near 6.46 h , but they are not at exactly the same values. This suggests one may be able to extract more detailed information on the frequency variation of Δ by using such narrow windows.

I then took a set of narrower windows of $\pm 200 \text{ Hz}$ centred on a frequency ν and moved the windows through the frequency range $1200\text{--}2500 \mu\text{Hz}$ in steps of $5 \mu\text{Hz}$, and determined the location of the peaks t_k in the autocorrelation power near 6.46 h , and hence a local value of the large separations as $\Delta(\nu) = 1/2t_k$. The results are shown in Fig. 4a. Superimposed on this curve are the values of the large separations (and their 1σ formal errors) as determined by Appourchaux et al. (2008). The agreement is poor. However there is considerable uncertainty in the determination of these frequencies: it is difficult to decide which modes are $\ell = 0, 2$ pairs and which are rotationally split $\ell = 1$ modes (see e.g. Kallinger et al. 2008), there is uncertainty in determining the rotation and angle of inclination of the star, and the values of frequencies determined by fitting pairs of modes, longer sets, or a complete set of $14n$ values do not agree. The values at high

frequency are particularly difficult to extract. This is illustrated in Fig. 4b which shows the results for 23 different frequency sets determined for HD 49933 from the initial run 60 day data set with different fitting assumptions and extraction procedures (Verner, private communication). The variation is large.

Fortunately HD 49933 has subsequently been observed during a 132 day run on CoRoT – leading to improved frequency determinations. These results have yet to be published by the CoRoT Data Analysis Team but are expected to give much better agreement with the variation of large separations obtained by frequency windowed autocorrelation.

Figure 4c shows the same results but for a narrow frequency window of full width $200 \mu\text{Hz}$. The fit to the mean of the several frequency sets is improved. The curve has a modulation period of $\approx 43 \mu\text{Hz}$, which is half of the mean value of the large separations. As shown below this is because the “*half large separations*”,

$$\Delta_{10}(n) = \nu_{n,1} - \nu_{n,0}, \quad \Delta_{01}(n) = \nu_{n,0} - \nu_{n-1,1}, \quad (3)$$

between neighbouring modes of $\ell = 0, 1$ are not expected to be equal and, with a narrow window, the resulting autocorrelation time depends on whether an $\ell = 0, 1$ or $\ell = 1, 0$ pair is in the central part of the window. (For the Sun the difference between Δ_{10} and Δ_{01} varies between $5\text{--}10 \mu\text{Hz}$.) The signal is also influenced by the $\ell = 0, 2$ small separations but may still provide a useful diagnostic of the internal structure of the star.

3. HD 175726: a low signal to noise example

HD 175726 is a F9/G0 ($m_v = 6.72$) star observed for 27 days in the first short run of CoRoT. Details of the properties of this star are given in Mosser et al. (2009). As for HD 49933, the level 2 CoRoT time series has a 32 s cadence, and ($\sim 9 \text{ min}$) gaps due to passage through the South Atlantic anomaly filled by linear interpolation. The major variations in the light curve were removed by subtracting off a running mean, and experiments using different running means and different gap-filling procedures showed that the results described below were largely independent of these procedures. The power spectrum is shown in Fig. 5a, individual p-modes cannot be seen in the spectrum, but in the boxcar spectrum shown in Fig. 5b a slight excess of power can be seen in the p-mode frequency range $1000\text{--}3000 \mu\text{Hz}$. It is therefore worthwhile to look for evidence of variations of the large separations with frequency using narrow-windowed autocorrelations – indeed it was the challenge presented by this star that initiated the research reported here. A detailed description of analysis techniques and efforts to extract frequencies are given in Mosser et al. (2009).

The autocorrelation power, filtered by a \sin^2 window between $1000\text{--}3000 \mu\text{Hz}$, is shown in Fig. 5c. The highest peak away from zero is at $t = 5.69 \text{ h}$, which, if due to the large separations, gives a mean large separation $\Delta = 1/2t = 97.7 \mu\text{Hz}$. Also shown in Fig. 5c as the dotted line is the autocorrelation power obtained using a boxcar power spectrum with averaging over $15 \mu\text{Hz}$. It should be stressed that one is here fighting to extract a signal from the noise and the results will differ somewhat depending on how one treats the data: filtering out low frequency variations, filling gaps by linear interpolation with or without added noise, and suppressing or not harmonics of the orbital period. Nevertheless the general behaviour is found with different procedures (see Mosser et al. 2009, for a more detailed discussion on the statistical robustness of the results).

I then took a set of narrow windows centred on frequencies ν in the range $1500\text{--}2500 \mu\text{Hz}$ with half width $\delta\nu = 300$,

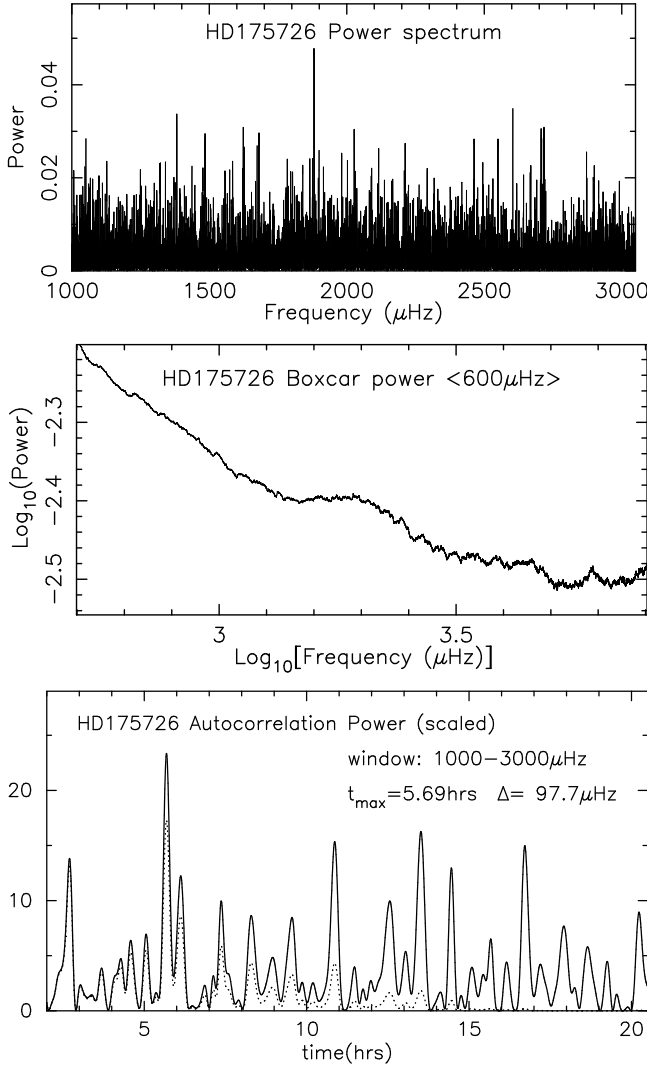


Fig. 5. **a)** Power spectra for HD 175726; **b)** result of applying a boxcar of $600 \mu\text{Hz}$ full width. **c)** Autocorrelation power for HD 175726 for very wide window of $\pm 1000 \mu\text{Hz}$. The dominant peak is at a correlation time of 5.69 h corresponding to a large separation of $97.7 \mu\text{Hz}$.

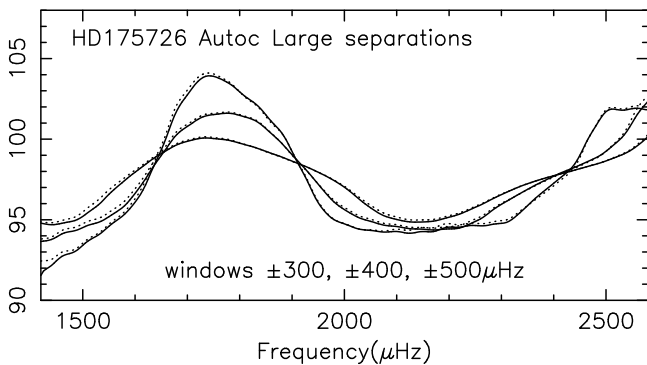


Fig. 6. Variation of Large separation with frequency for HD 175726 for a set of narrow window of $\pm 300, 400, 500 \mu\text{Hz}$.

$400, 500 \mu\text{Hz}$ and identified the peak near 5.7 h and hence a local value of $\Delta(\nu)$. With a mean large separation of $\approx 98 \mu\text{Hz}$ the windows were wide enough to ensure that (if they exist) there are several $\ell = 0, 1$ modes within the window (6–10) which should suppress the variation between pairs of modes and give a local average of the large separation. The results are shown in

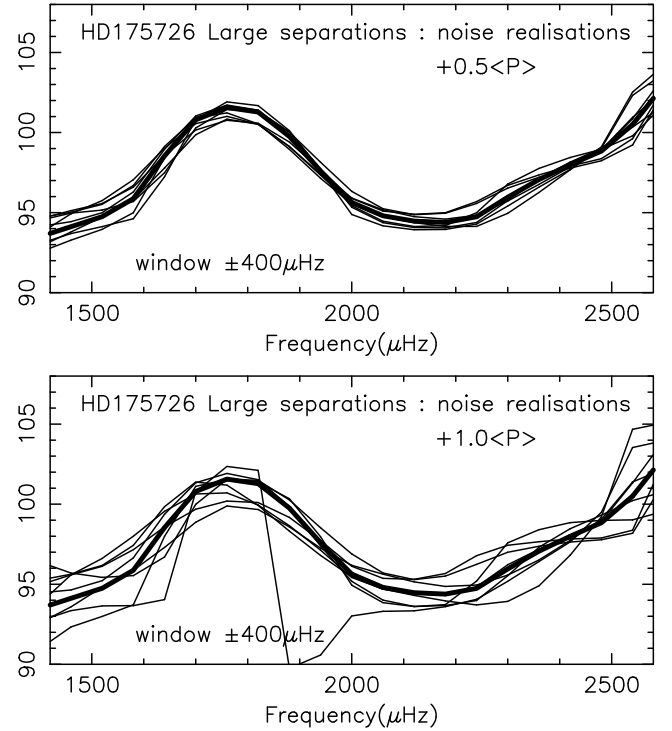


Fig. 7. Autocorrelation power for HD 175726 \sin^2 windowed with half width $400 \mu\text{Hz}$, with exponential noise added to the power spectrum: **a)** $0.5 \langle P \rangle$, **b)** $1.0 \langle P \rangle$.

Fig. 6; all 3 curves show a similar quasi-periodic behaviour but the smaller the window the larger the amplitude of the variation. This is to be expected since the larger windows span a wider range around the peaks and troughs and therefore reduce the amplitude of the quasi-periodic variation. Such a quasi-periodic behaviour is caused by the variation in the surface phase shift $\alpha(\nu)$ due to the HeII ionisation zone, (Vorontsov & Zharkov 1988,9; Brodskii & Vorontsov 1989), and also by the internal phase shifts $\delta_\ell(\nu)$ (Roxburgh & Vorontsov 1994; Roxburgh 2009). Whilst it is admittedly difficult to extract any such periodicity from Fig. 6 there is some indication of a period $\sim 1000 \mu\text{Hz}$, which were it due to the HeII ionisation layer, corresponds to an acoustic depth of ~ 500 s. This is not unreasonable for such a star, but the amplitude is larger than theoretical expectations.

I then added exponentially distributed noise to the power spectrum with a mean of 0.5 and 1.0 times the average power $\langle P \rangle$ in the frequency range $1500\text{--}2500 \mu\text{Hz}$, and repeated the calculations for a $\pm 400 \mu\text{Hz}$ window; the results are shown in Figs. 7a and b, the thick line being the result with no added noise. These results suggest that the signal is there, but this does not of course prove that the signal is due to the variation $\Delta(\nu)$. Estimates of the error on the autocorrelation due to time resolution, window width, and interference between signal and noise are considered in Mosser et al. (2009). They also test the reliability of the shape displayed in Fig. 6 with an H0 test and concluded that the hypothesis that the signal is real is only rejected at the $\sim 1\text{--}10\%$ level.

As mentioned above, one is here fighting against the noise – just how much is illustrated in Fig. 8, which gives the autocorrelation power for a window of half width $350 \mu\text{Hz}$ centred on frequencies in the range 1500 and $2300 \mu\text{Hz}$. A peak around $98 \mu\text{Hz}$ can just be seen, although how significant it is needs to be the subject of further investigation. The dotted line in Fig. 8

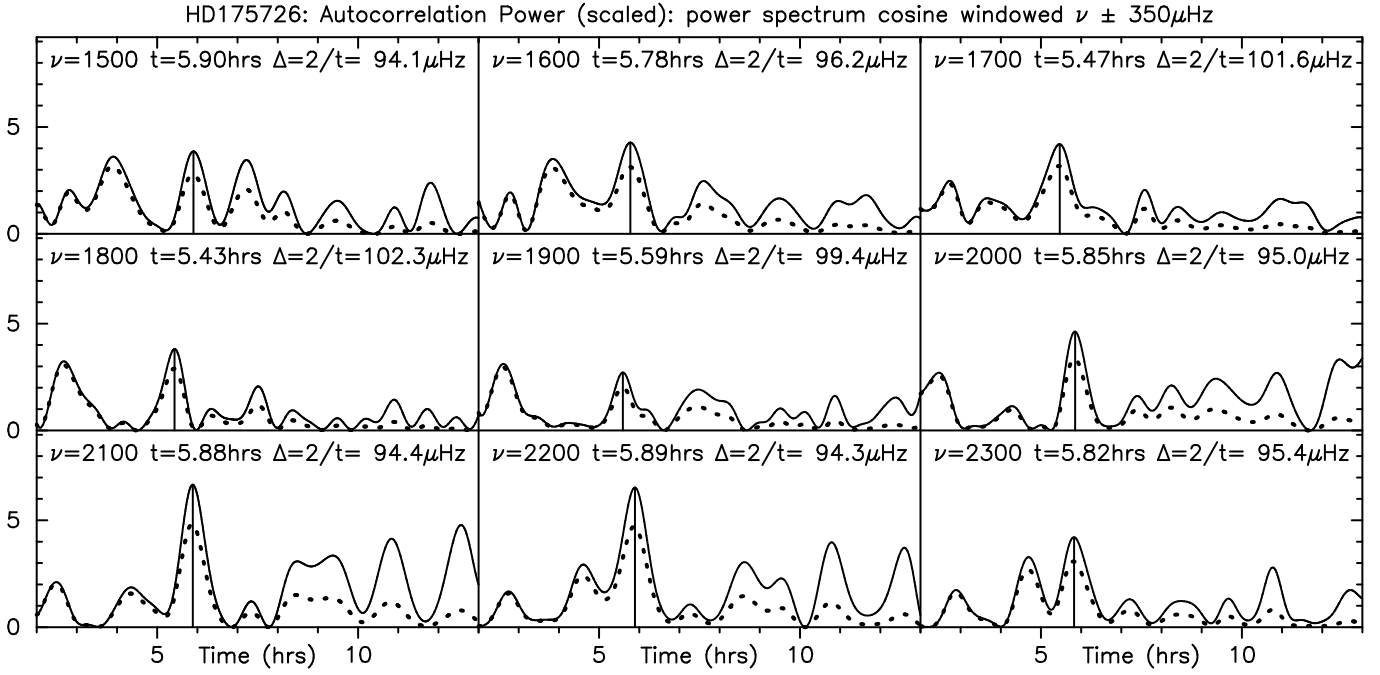


Fig. 8. Autocorrelation power for HD 175726 \sin^2 windowed with half width $350 \mu\text{Hz}$ window. The dotted curves are obtained from a $15 \mu\text{Hz}$ boxcar spectrum.

is the autocorrelation power using a boxcar spectrum averaging over $15 \mu\text{Hz}$.

4. Analysis and modelling

I first take the simplest possible model where the peaks in the frequency power spectrum are given by delta functions so

$$P(\nu) = \sum A_i \delta(\nu - \nu_i), \quad (4)$$

where ν_i are the eigenfrequencies. For simplicity of presentation I confine my attention to the case where the only modes are those with $\ell = 0, 1$ so that $\{\nu_i\} = \nu_{n,0}, \nu_{n,1}, \nu_{n+1,0}, \nu_{n+1,1}, \dots$

As shown by Roxburgh and Vorontsov (2000), by matching the inner and outer solutions of the equations governing the oscillations at an intermediate radius r_f , the eigenfrequencies $\nu_{n,\ell}$ satisfy the eigenfrequency equation

$$\nu_{n,\ell} = \Delta(n + \ell/2) + \frac{\Delta}{\pi}(\alpha(\nu) - \delta_\ell(\nu)), \quad \Delta = \frac{1}{2T}, \quad (5)$$

where $\alpha(\nu)$ is an ℓ independent surface phase shift due to the structure of the outer layers, $\delta_\ell(\nu)$ the ℓ dependent internal phase shifts due to the structure of the inner layers, and T is the acoustic radius of the star. The values of α, δ_ℓ are almost independent of the location of the matching radius for $0.8 \leq r_f/R \leq 0.98$. (Roxburgh & Vorontsov 2000; Roxburgh 2009). Note that Δ is not exactly the large separation, in fact

$$\Delta_{n,\ell} = \Delta \left(1 + \frac{1}{\pi} [\alpha(\nu_{n,\ell}) - \alpha(\nu_{n-1,\ell}) - \delta_\ell(\nu_{n,\ell}) + \delta_\ell(\nu_{n-1,\ell})] \right). \quad (6)$$

Consider first the simple model consisting solely of a pair of adjacent modes $\nu_0 = \nu_{n,0}$ and $\nu_1 = \nu_{n,1}$. The Fourier transform of the power spectrum consisting of these two modes is

$$F(t) = \int_0^\infty [a_0 \delta(\nu - \nu_0) + a_1 \delta(\nu - \nu_1)] e^{2\pi i \nu t} dt = e^{2\pi i \nu_0 t} [a_0 + a_1 e^{2\pi i (\nu_1 - \nu_0) t}], \quad (7)$$

where the a_i are the amplitudes of the modes in the power spectrum.

The autocorrelation power (the power spectrum of the power spectrum) is $|F|^2$ which is then

$$A^2 = a_0^2 + a_1^2 + 2a_0 a_1 \cos[2\pi(\nu_1 - \nu_0)t]. \quad (8)$$

This has its first peak (beyond zero) when $(\nu_1 - \nu_0)t = 1$. If $\alpha, \delta_0 = \delta_1$ are all constant then from Eq. (5) $\nu_1 - \nu_0 = \Delta/2$ and hence the first peak occurs at $t = 4T$. If the amplitudes are the same for a sets of pairs of modes then they all add up to give the same position of the first maximum (see Roxburgh & Vorontsov 2006a).

However if, as is the real case, $\alpha, \delta_0, \delta_1$ vary with frequency and $\delta_1 \neq \delta_0$, the above pair analysis remains valid but $\nu_1 - \nu_0 \neq \Delta/2$. I therefore define the *half large separations*

$$\Delta_{10}(n) = \nu_{n,1} - \nu_{n,0}, \quad \Delta_{01}(n) = \nu_{n,0} - \nu_{n-1,1}. \quad (9)$$

Their sum gives the ordinary large separations

$$\Delta_0(n) = \nu_{n,0} - \nu_{n-1,0} = \Delta_{01}(n) + \Delta_{10}(n-1), \quad (10)$$

and difference the small separations (Roxburgh 1993, 2009)

$$d_{01}(n) = \nu_{n,0} - (\nu_{n,1} + \nu_{n-1,1})/2 = (\Delta_{01}(n) - \Delta_{10}(n-1))/2. \quad (11)$$

The power from the pair of modes $\{\nu_{n,0}, \nu_{n,1}\}$ is

$$A^2 = (a_0^2 + a_1^2 + 2a_0 a_1 \cos[2\pi \Delta_{10}(n)t]), \quad (12)$$

which has its first maximum at $t = 1/\Delta_{10}(n)$, whereas that from the pair $\{\nu_{n,1}, \nu_{n+1,0}\}$ is

$$A^2 = (a_0^2 + a_1^2 + 2a_0 a_1 \cos[2\pi \Delta_{01}(n+1)t]), \quad (13)$$

and has its first maximum at $1/\Delta_{01}(n+1)$.

As can be seen from Eq. (5)

$$\Delta_{10}(n) = \Delta \left(\frac{1}{2} + \frac{\alpha(\nu_{n,1}) - \alpha(\nu_{n,0})}{\pi} - \frac{\delta_1(\nu_{n,1}) - \delta_0(\nu_{n,0})}{\pi} \right) \quad (14)$$

$$\Delta_{0,1}(n) = \Delta \left(\frac{1}{2} + \frac{\alpha(\nu_{n,0}) - \alpha(\nu_{n-1,1})}{\pi} - \frac{\delta_0(\nu_{n,0}) - \delta_1(\nu_{n-1,1})}{\pi} \right) \quad (15)$$

so even if $\alpha, \delta_0, \delta_1$ constant, but $\delta_1 \neq \delta_0$, these differ by

$$\Delta_{01}(n+1) - \Delta_{10}(n) = \frac{2\Delta}{\pi}(\delta_1 - \delta_0). \quad (16)$$

For the Sun the difference $\Delta_{01} - \Delta_{10}$ varies with frequency ν from about $10 \mu\text{Hz}$ at $\nu = 1000 \mu\text{Hz}$ to $5 \mu\text{Hz}$ at $\nu = 4000 \mu\text{Hz}$.

For a large set of eigenfrequencies within a window (including $\ell = 2, 3$) the situation is more complicated since one cannot just add up the power in each pair but have to take the full Fourier transform. However one can see that if there is only one pair $\{\nu_{n,0}, \nu_{n,1}\}$ in a window, the position of the first maximum will be different from that with just the overlapping pair $\{\nu_{n,1}, \nu_{n+1,0}\}$. If there are a number of such sets within a window then one may expect the first peak in the autocorrelation power spectrum to be determined by the full large separations Δ_0, Δ_1 . Note that, at least for the Sun, the difference between Δ_0 and Δ_1 is much smaller than that between Δ_{01} and Δ_{10} since the inner phase shifts δ_0, δ_1 differ by much more than the change in α, δ_0 , and δ_1 between adjacent frequencies. Of course the actual windowed autocorrelation for a small set of frequencies depends on the amplitudes of all significant modes ($\ell = 0, 1, 2, 3$) within the window and on noise peaks. For a theoretical model this could be calculated but for a real data set one can do no more than predict that for very narrow windows the first peak of the autocorrelation function will vary depending on whether it is an $\ell = 0, 1$ pair or an $\ell = 1, 0$ pair at the centre of the window, but for a wider window the peak is determined by a locally averaged large separation. This is the origin of the oscillatory behaviour of the $\pm 100 \mu\text{Hz}$ windowed results for HD 49933 displayed in Fig. 4c. It offers the possibility of determining the inner phase shift difference $\delta_1 - \delta_0$ as a function of frequency, which is an important diagnostic of the stellar interior and convective boundaries (Roxburgh 2009).

To demonstrate that this technique can, in principle, work I constructed a theoretician's ideal artificial noise free power spectrum by prescribing a surface phase shift $\alpha(\nu)$ and inner phase shifts $\delta_\ell(\nu)$, with the eigenfrequencies determined by the *Eigenfrequency equation* (Roxburgh & Vorontsov 2000)

$$2\pi\nu_{n,\ell}T = (n + \ell/2)\pi + \alpha(\nu) - \delta_\ell(\nu). \quad (17)$$

To produce a model with similar characteristics to HD 175726 the model had an acoustic radius of $T = 5000$ s, and α and δ_ℓ given by

$$\alpha(\nu) = \frac{4.5}{0.5 + \nu_2} \left(1 + \frac{0.03375}{\nu_2^2} \sin[2\pi\nu/\nu_0 + 1.5] \right) \quad (18)$$

$$\delta_\ell(\nu) = \frac{1}{\nu_2} [0.7 + 0.05\ell(\ell + 1)], \quad (19)$$

where $\nu_2 = \nu/2000 \mu\text{Hz}$ and $\nu_0 = 650 \mu\text{Hz}$.

The individual line profiles were taken to be Lorentzians of width $4 \mu\text{Hz}$ and the amplitude ratios for modes with $\ell = 0, 1, 2, 3$ in the ratios 1, 1.334, 0.35, 0.016. The resulting powers spectrum was then scaled by the factor $\exp[-\lambda(\nu_2 - 1)^2]$ with $\lambda = 40/3$. The resulting model power spectrum is shown in Fig. 9a. Applying the narrow windowed autocorrelation analysis with windows of $\pm 150, 300, 500 \mu\text{Hz}$ yielded the curves displayed in Fig. 9b, together with the values of Δ derived from the

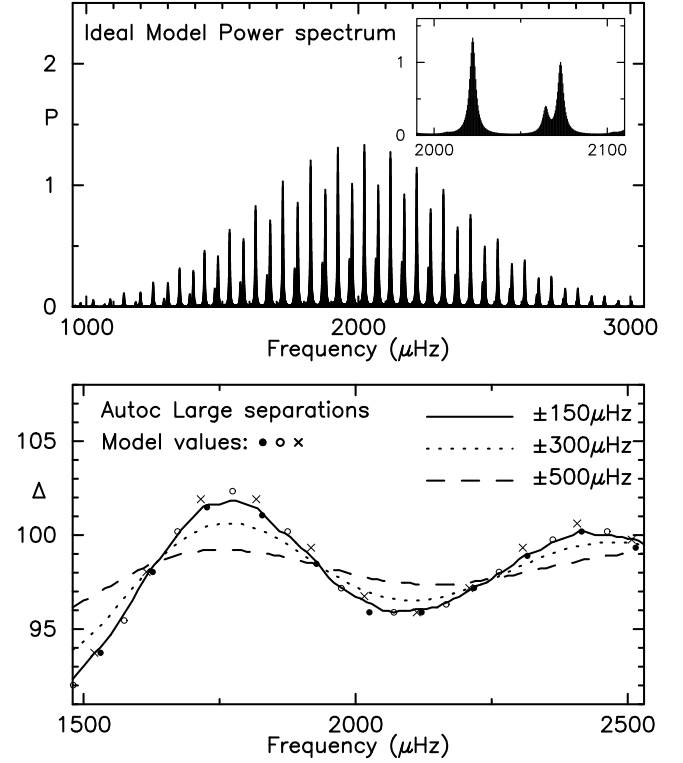


Fig. 9. a) Ideal model power spectrum; the inset shows an example of the line profiles. b) Large separation $\Delta(\nu)$ determined from frequency windowed autocorrelations. The model values of Δ_ℓ are also shown.

frequencies of the model. As expected the wider windows depart further from the actual large separations since the wider the window the greater the smoothing of the actual separations. The narrowest window shows the beginnings of a periodic modulation due to the difference between the *half large separations* Δ_{01} and Δ_{10} .

I then added exponential noise to this artificial power spectrum as shown in Fig. 10a, which is on the same scale as Fig. 9a. The noiseless power spectrum is shown in white in this figure. The variation of the large separations with frequency as determined from frequency windowed autocorrelation with windows of $\pm 200, 300, 500 \mu\text{Hz}$ are shown in Fig. 10b.

Comparison of these results with those for HD 175726 shown in Fig. 6 suggests one can extract some information on the large scale variations of $\Delta(\nu)$ by this procedure.

Returning to the noiseless model power spectrum I then took very narrow frequency windows of $\pm 60, 80, 100 \mu\text{Hz}$ to test whether this can give the *half large separations* Δ_{01}, Δ_{10} defined above in Eq. (9). Figure 11a shows the results for the full model power spectrum. The autocorrelation values show the same general behaviour as the exact values of the frequencies but somewhat offset. That this is due to the contribution of the $\ell = 2$ modes is clear from Fig. 11b where the model power spectrum was taken to only have $\ell = 0, 1$ modes; in this case the *half large separations* are faithfully reproduced.

5. HD 181420

HD 181420 is an F2 ($m_v = 6.57$) star observed by CoRoT in a 156 day run (Michel et al. 2008, Barban et al. 2009). Details of the data reduction are given in Barban et al. (2009), but are essentially the same as described above for HD 49933. The power spectrum is given in Fig. 12a and a boxcar spectrum in

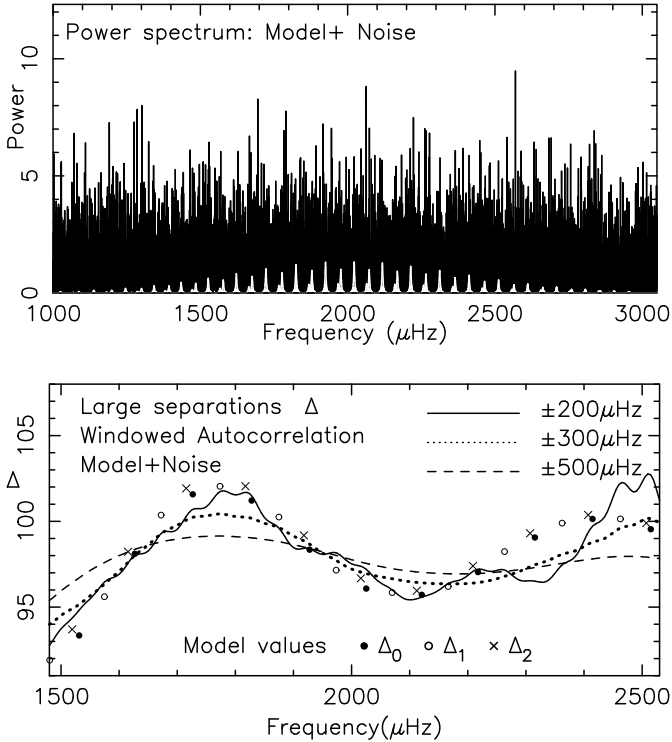


Fig. 10. **a)** Ideal model power spectrum + added noise; the noiseless power spectrum is shown in white. **b)** Large separation $\Delta(\nu)$ determined from frequency windowed autocorrelations of the noisy spectrum.

Fig. 12b. There is excess power in the p-mode frequency range 1200–2000 μHz . Details of the extraction of frequencies are given in Barban et al. (2009) but as with HD 49933 it is difficult to discriminate between peaks corresponding to modes $\ell = 0, 2$ and those corresponding to $\ell = 1$; Barban et al. (2009) give frequency sets for both (Scenarios 1 and 2). I here give the results of applying narrow frequency windowed autocorrelations to the time series and see if one can differentiate between them.

Figure 13a shows the autocorrelation power in a wide (1200 μHz) window with a clean peak at $t_m = 7.37$ hrs corresponding to a mean large separation of 75.4 μHz , and a secondary peak at $2t_m$. The rapid decrease in amplitude between the two peaks is again indicative of the large line widths in the power spectrum. Figure 13b shows the results of using a narrower frequency window of full width 400 μHz . Again there is variation of Δ with frequency.

I then used narrow windows of $\pm 100, 150$ μHz and determined the local large separations $\Delta(\nu)$, the results are in Fig. 14a for Barban’s Scenario 1, and in Fig. 14b for Scenario 2. The points and error bars are from the frequency sets given in Barban et al. There is not much to choose between the fit of the ± 100 μHz windowed results to the separations determined from the frequencies – possibly Scenario 2 is a little better than Scenario 1.

Since the values of Δ vary substantially with frequency a better test would be to use even narrower windows which give the half large separations Δ_{01}, Δ_{10} as in Fig. 11. The results for a window of ± 67 μHz are shown in Fig. 15 for both scenarios. Here Scenario 2 gives a better fit than Scenario 1, suggesting that Scenario 2 is the correct identification of the modes.

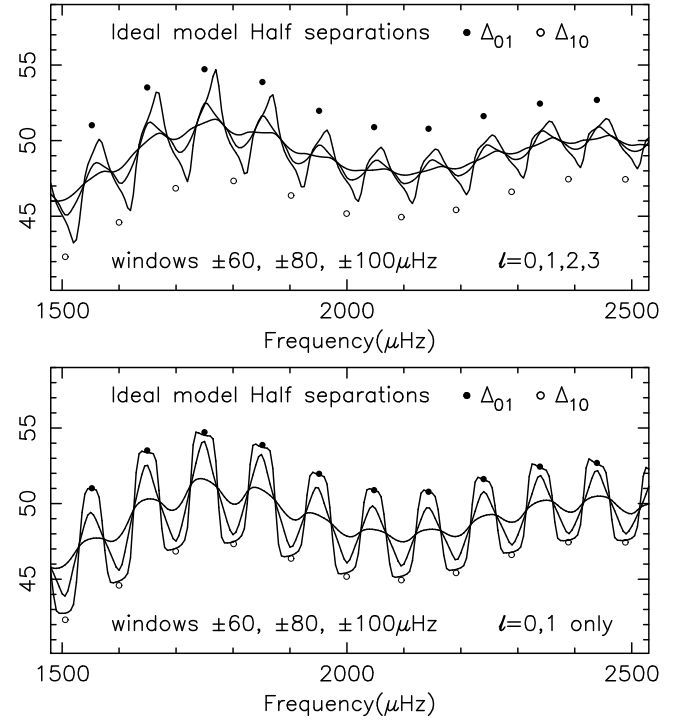


Fig. 11. **a)** Half large separation $\Delta_{01}(\nu), \Delta_{10}(\nu)$ determined from very narrow windowed autocorrelations of the full noiseless power spectrum with modes $\ell = 0, 1, 2, 3$. The points are the values from the frequencies of the model. **b)** Half large separation $\Delta_{01}(\nu), \Delta_{10}(\nu)$ determined from very narrow windowed autocorrelations of the power spectrum with only modes with $\ell = 0, 1$.

6. HD 181906

HD 181906 is an F8 ($m_v = 7.65$) star observed by CoRoT for 156 days, that displays very low p-mode power in the frequency range 1000–2500 μHz (Michel et al. 2008, García et al. 2009); it is considerably fainter than HD 49933 and HD 181420. Figure 16a gives the power spectrum and Fig. 16b the boxcar spectrum which shows an excess of power in the p-mode frequency range 1100–2300 μHz . The power spectrum was derived from the time series by the same procedures as for HD 49933. Figure 16c gives the autocorrelation power for a wide frequency window of 1200 μHz . There is a clear peak at $t_m = 6.45$ hrs corresponding to a mean large separation of 85.1 μHz and a small second peak at $2t_m$.

Extracting frequencies from the power spectrum is difficult as the signal to noise is small and rotational splitting, line widths and ($\ell = 0, 2$) separations all overlap. By making severe constraints on the fitting procedure (mode widths and mode heights set equal over the fitted frequency range) values of nine $\ell = 0, 1, 2$ frequencies were obtained by García et al. (2009). These differ from those obtained for a smaller frequency sets (5 n values) obtained by Verner (private communication) with less rigid constraints on the fitting procedure, and are only marginally consistent within the larger formal errors of Verner’s analysis. Again there is uncertainty over which peaks correspond to $\ell = 0, 2$ and which to rotationally split $\ell = 1$ modes; these two alternatives are labelled Scenario A and B in García et al.

I then used narrow windows of $\pm 100, 150$ μHz and determined the local large separations $\Delta(\nu)$ from the peaks in the autocorrelation power; the results are in Fig. 17a for García’s Scenario A and Fig. 17b for Scenario B. The circles in these diagrams are the values obtained with García’s frequencies and the

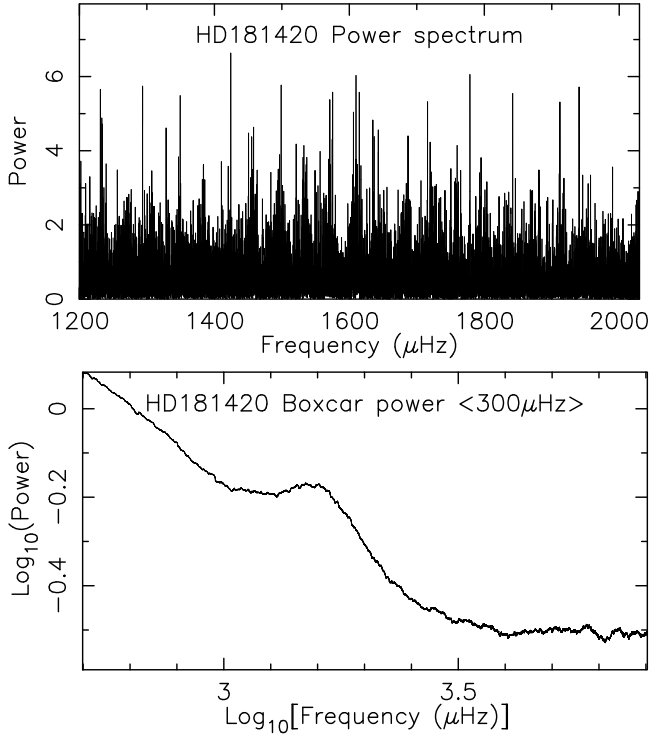


Fig. 12. a) HD 181420 Power spectrum. b) Boxcar power spectrum full width 300 μHz .

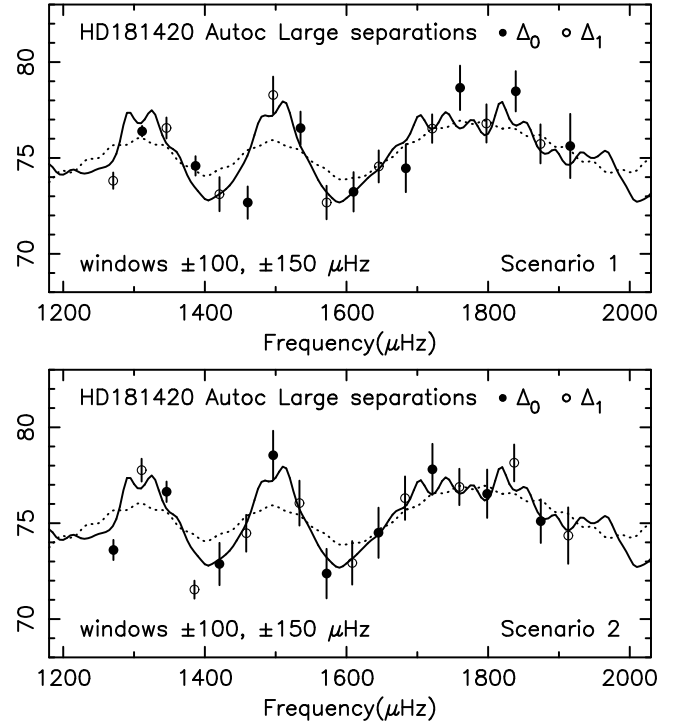


Fig. 14. Autocorrelation Large separations for HD 181420 with $\pm 100, 150 \mu\text{Hz}$ windows. a) Scenario 1. b) Scenario 2.

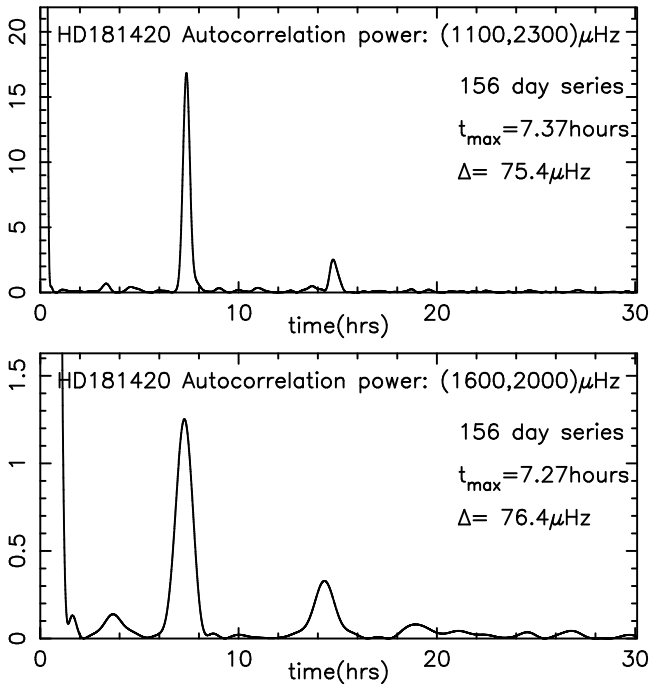


Fig. 13. a) Autocorrelation power for HD 181420 with 1200 μHz window. b) Same as a) but for narrower window of 400 μHz .

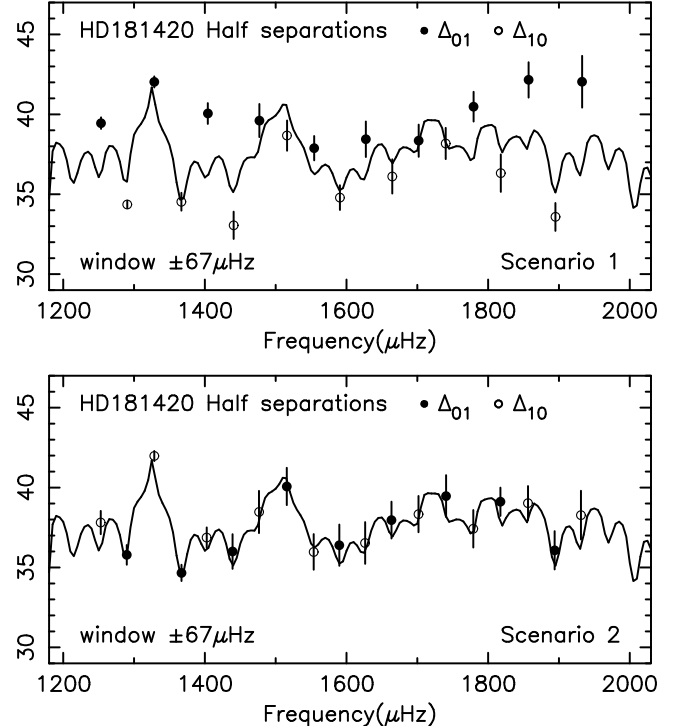


Fig. 15. Autocorrelation Half Large separations for HD 181420 with $\pm 67 \mu\text{Hz}$ window. a) Scenario 1. b) Scenario 2 (Barban et al. 2009).

crosses those from 3 different analyses by Verner, each with their estimated formal errors. The considerable difference between values from different frequency extraction algorithms illustrates the difficulty in deriving reliable values of the frequencies for this star. Figs. 18 gives the values of the half large separations using an $\pm 80 \mu\text{Hz}$ window for the two Scenarios. The agreement between the autocorrelation values and those from any frequency

set is not good. Further work needs to be done on both frequency extraction and the windowed autocorrelations to see if one can obtain agreement similar to that for HD 181420 in Fig. 15b.

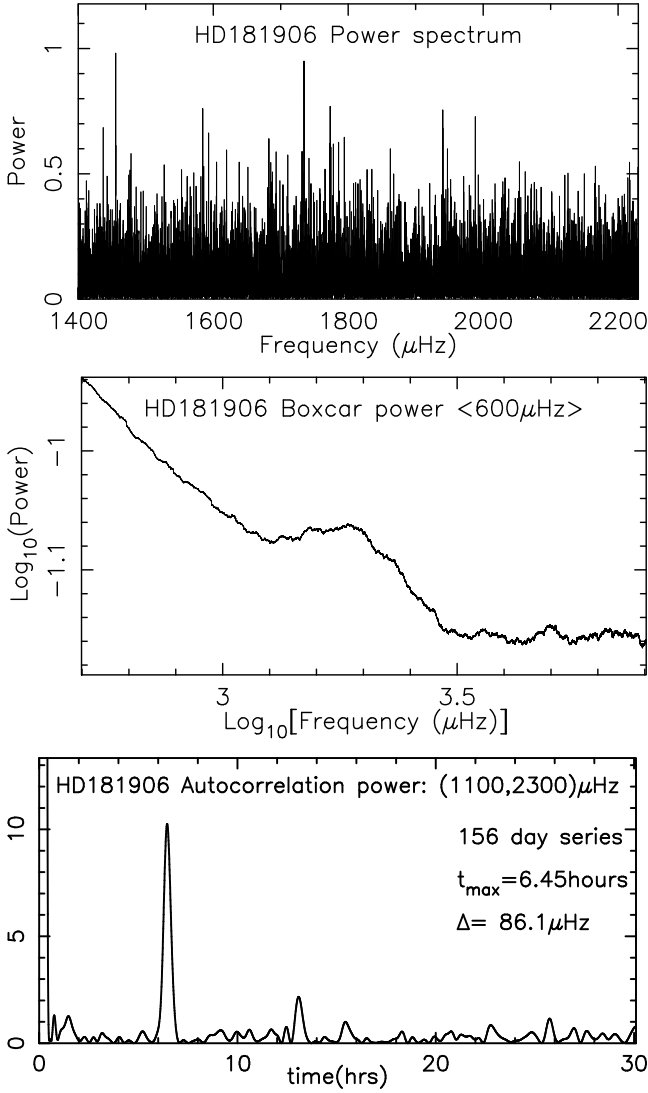


Fig. 16. a) Power spectrum of HD 181906. b) Boxcar power spectrum with 600 μHz average. c) Autocorrelation power with 1200 μHz window.

7. Conclusions

The principal goal of this paper was to show that narrow frequency windowed autocorrelation can, in principle, reveal information on the variation of the large separation with frequency and therefore constitutes a tool that might be useful in obtaining some information about a star even when individual frequencies cannot be extracted, or modes cannot be identified. Much remains to be done to refine the technique: the theoretical analysis needs to be further developed and the nature of the interaction of noise with the autocorrelation power better understood. However the analysis presented here suggests that narrow frequency windowed autocorrelations can yield the variation with frequency of the large separations $\Delta(\nu)$ and that very narrow windows can yield the half large separations Δ_{01} and Δ_{10} and thereby help to resolve the uncertainty over mode recognition between $\ell = 0, 2$ and $\ell = 1$ modes.

Since the difference between the half large separations is determined by the inner phase shift difference $\delta_1 - \delta_0$, this gives a diagnostic of the internal structure of a star (Roxburgh 2009). This will be explored in a subsequent communication.

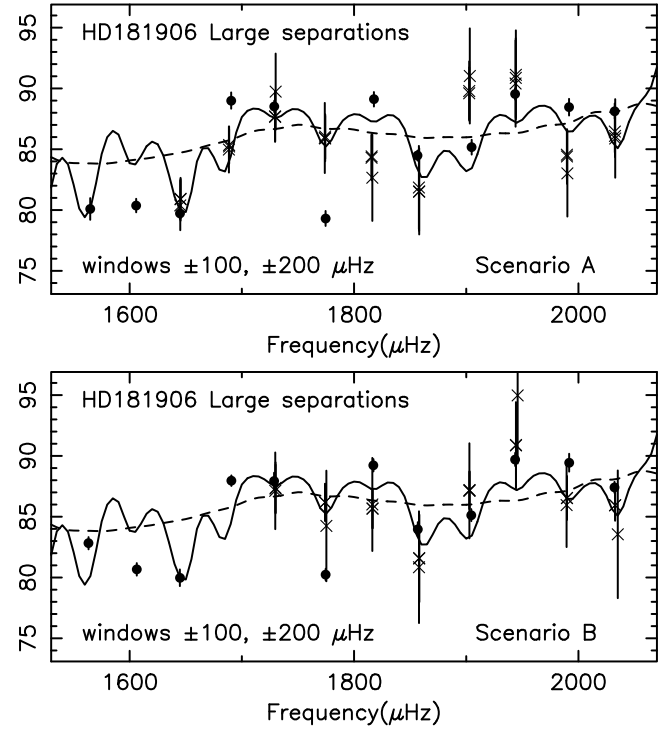


Fig. 17. a) Autocorrelation Large separations for HD 181906 with $\pm 100, 150 \mu\text{Hz}$ windows Scenario A. b) as above but for scenario B.

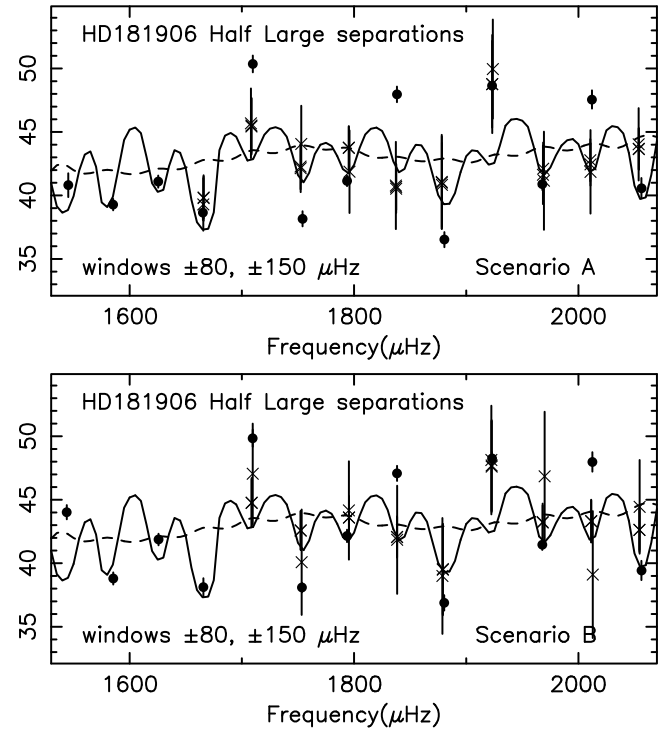


Fig. 18. a) Autocorrelation Half Large separations for HD 181906 with $\pm 80 \mu\text{Hz}$ window for Scenario A. b), as above but for Scenario B.

Acknowledgements. I thank the UK Science and Technology Facilities Council (STFC) which supported this work under grants PPA/G/S/2003/00137 and PP/E001793/1 and Graham Verner for the determination of the unpublished frequencies used in this paper.

References

- Appourchaux, T., Michel, E., Auvergne, M., et al. 2008, *A&A*, 488, 705
- Barban, C., Deheuvels, S., Baudin, F., et al. 2009, *A&A*, 506, 51
- Brodskii, M. A., & Vorontsov, S. V. 1989, *Astron. Zh.*, 15, 61 (English version: *Sov. Ast. Lett.* 1989, 15, 27)
- Gabriel, M., Grec, G., Renaud, C., et al. 1998, *A&A*, 338, 1109
- García, R., Régulo, C., Samadi, R., et al. 2009, *A&A*, 506, 41
- Kallinger, T., Gruberbauer, M., Guenther, D. B., Fossati, L., & Weiss, W. W. 2008, *A&A*, submitted [[arXiv:0811.4686v1](https://arxiv.org/abs/0811.4686v1)]
- Kholikov S., & Hill, F. 2008, *Sol. Phys.*, 251, 157
- Michel, E., Baglin, A. Auvergne, M., et al. 2008, *Science*, 322, 558
- Mosser, B., Michel, E., Appourchaux, T., et al. 2009, *A&A*, 506, 33
- Press, W. H., Flannery, B. P., Teukolsky A. S., & Vetterling, W. T. 1992, *Numerical Recipes (CUP)*, 383
- Roxburgh, I. W. 1993, in *PRISMA, Report of Phase A Study*, ed. T. Appourchaux, T., *ESA SCI(93)*, 31
- Roxburgh, I. W. 2009, *A&A*, 493, 185
- Roxburgh, I. W., & Vorontsov, S. V. 1994, *MNRAS*, 267, 297
- Roxburgh, I. W., & Vorontsov, S. V. 2000, *MNRAS*, 317, 141
- Roxburgh, I. W., & Vorontsov, S. V. 2006a, *MNRAS*, 369, 1491
- Roxburgh, I. W., & Vorontsov, S. V. 2006b, *MNRAS*, 379, 801
- Samadi, R., Fialho, F., Costa, J. E. S., et al. 2006, *ESA SP 1306*, 317, corrected in [[arXiv:astro-ph/0703354](https://arxiv.org/abs/astro-ph/0703354)]
- Vorontsov, S. V., & Zharkov, V. N. 1988, *Itogi Nauki Tek, Ser Astro*, 38, 253 (English version: *Astro & Space Phys Rev.* 1989, 7, 1)
- Vorontsov, S. V., & Zharkov, V. N. 1989, *Astro & Space Phys. Rev.*, 7, 1