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LETTER TO THE EDITOR

[Ar III]/[O III] and [S III]/[O III]: well-behaved oxygen abundance indicators for HII regions and star forming galaxies

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ABSTRACT

We propose two statistical methods to derive oxygen abundances in HII regions and star forming galaxies and calibrate them with a sample of several hundred giant HII regions in spiral and blue compact galaxies as well as of galaxies from the Sloan Digital Sky Survey. We show the advantages of our new abundance indicators over previous ones.

Key words. galaxies: abundances – galaxies: ISM – ISM: abundances – ISM: HII regions

1. Introduction

Emission lines due to photoionization of nebulae by massive stars are the most powerful indicators of the chemical evolution of galaxies in both the near and the intermediate redshift universe. The wealth of new data coming from large/deep galaxy spectroscopic surveys and from observations with integral field units on nearby galaxies should allow a detailed inventory of the chemical composition of star forming galaxies.

Abundance determinations in HII regions are relatively straightforward if the electron temperature, T_e , can be measured directly from the observations (although strong biases may affect abundance derivations in high metallicity objects, even if lines used for T_e measurements are observed, see Stasińska 2005). In many cases, however, T_e diagnostic lines are too faint to be detected. Pagel et al. (1979) and Alloin et al. (1979) pioneered methods that allow one to estimate the metallicity¹ using strong lines only. These methods assume that all HII regions are essentially characterized by their metallicity, other parameters such as the hardness of the ionizing radiation and the ionization parameter being actually linked to metallicity. Empirically, this proves true, at least statistically, as has been shown by many theoretical and observational studies. The biggest problem is the translation of an observed line ratio into metallicity. We will not review here the large amount of literature devoted to this subject. References can be found in Pérez-Montero & Díaz (2005), Pettini & Pagel (2004, hereafter PP04) and Stasińska (2004). There are two important issues. One is the choice of the metallicity indicator, the other is the proper abundance calibration. Here, we mostly discuss the first of the two, as we believe that time is not ripe yet to discuss the second: one really must understand the physics of metal-rich HII regions to be confident in any calibration at high metallicity. This requires the systematic use of mid-infrared data to efficiently constrain the modelling of metal-rich HII regions.

The most widely used indexes are the following: O_{23} , i.e. $([O III] \lambda 4959, 5007 + [O II] \lambda 3727)/H\beta$, introduced by Pagel et al. (1979), then refined by Mc Gaugh (1991) and Pilyugin (2001a);

O_3N_2 , i.e. $[O III] \lambda 5007/[N II] \lambda 6584$ introduced by Alloin et al. (1979) and reconsidered by PP04; N_2 , i.e. $[N II] \lambda 6584/H\alpha$ first proposed by Storchi-Bergmann et al. (1994) and Van Zee et al. (1998), and last recalibrated by PP04. Another index that is becoming popular, although it requires observations in the far red, is S_{23} , i.e. $([S III] \lambda 9069 + [S II] \lambda 6716, 6731)/H\alpha$ (Vílchez & Esteban 1996; Pérez-Montero & Díaz 2005). The main problem of the O_{23} index is that it is double valued. Therefore another index is needed to decide whether one has to chose the low-metallicity or the high-metallicity solution. This is not practical if dealing with data sets where both high and low metallicities are expected to be present. The S_{23} index is better since it is monotonic up to about solar metallicity, but a bend at higher metallicities is expected. The O_3N_2 and N_2 indexes, on the other hand, have the virtue of being single valued. However, the reasons for this are not purely “physical” but are partly due to “astrophysical” causes: it is an observed fact that, as metallicity increases, both the hardness of the ionizing radiation and the ionization parameter of HII regions tend to decrease, and that the N/O ratio increases on average (at least at high metallicities). The important role of the “astrophysical factors”, especially the variation of the N/O ratio should invite us to some caution when relying on them to derive metallicities of objects that might have a different nature and history than the objects used to calibrate the relation. Such might be the case of high redshift galaxies as compared to local dwarf galaxies.

Another problem that affects both the O_3N_2 and N_2 indexes is that they use a “low excitation line”, $[N II] \lambda 6584$ which may arise not only in bona fide HII regions, but also in the diffuse ionized medium. It is estimated that the diffuse ionized medium contributes up to one half of the total $H\alpha$ emission in galaxies (Thilker et al. 2002). Therefore, calibrations of the O_3N_2 and N_2 indexes based on giant HII regions might not be relevant for the interpretation of integrated spectra of galaxies.

Ideally, one would require from a metallicity index to: i) be single valued with respect to metallicity; ii) have a behaviour dominated by a well understood “physical” reason; iii) be unaffected by the presence of diffuse ionized gas; and iv) be independent of chemical evolution. We propose

¹ Throughout the paper the word “metallicity” is used with the meaning of “oxygen abundance”.

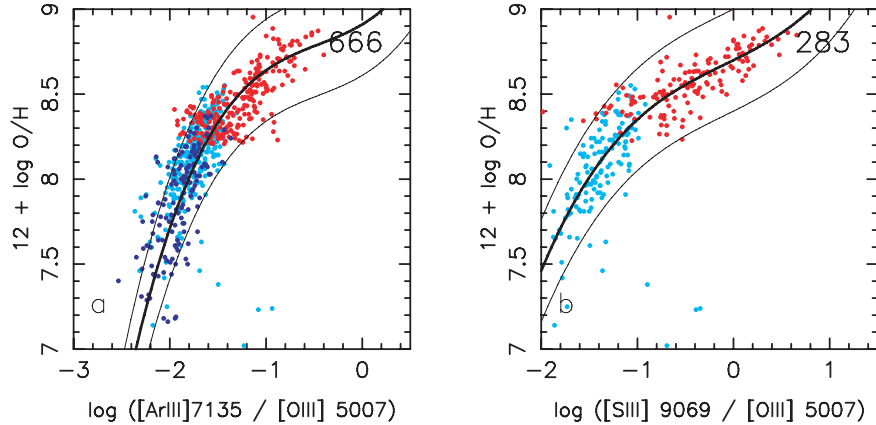


Fig. 1. **a)** $12 + \log O/H$ as a function of $\log Ar_3O_3$. Red: HII regions from the spiral galaxies sample; blue: HeBCD sample; cyan: SDSS DR3 sample. The thick curve represents our calibration defined by Eq. (1). The two thin curves are vertically displaced from the thick one by ± 0.3 dex. **b)** Same for $12 + \log O/H$ as a function of $\log S_3O_3$. The thick curve represents our calibration defined by Eq. (2).

two indexes: A_3O_3 , i.e. $[Ar III] \lambda 7135/[O III] \lambda 5007$ and S_3O_3 , i.e. $[S III] \lambda 9069/[O III] \lambda 5007$ that fulfill these requirements. In Sect. 2 we describe the data base sustaining this work, in Sect. 3, we present our new metallicity calibrators, and in Sect. 4, we discuss their merits with respect to the usual ones.

2. Data samples

We have built a large – though not exhaustive – data base of published reddening-corrected line intensities. Data for HII regions in spiral galaxies have been taken from Garnett & Kennicutt (1994), Garnett et al. (1997), van Zee et al. (1998), Bresolin et al. (1999, 2004, 2005), and Kennicutt et al. (2003). We have also used the sample of blue compact galaxies, observed by Izotov and coworkers for the quest of the pregalactic helium abundance, and referred to as the HeBCD sample in Izotov et al. (2006). The spectra in this sample have very good signal-to-noise ratios. Unfortunately, the wavelength range does not reach the $[S III] \lambda 9069$ line. Finally, we have used data from the SDSS DR3 sample, i.e. data from the 3rd data release of the Sloan Digital Sky Survey (Abazajian et al. 2005) that have been selected by Izotov et al. (2006) and are listed in their Table 1.

For both the HeBCD and SDSS DR3 samples, the oxygen abundances were computed using the temperature derived from $[O III] \lambda 4363/5007$ and following the prescriptions given in Izotov et al. (2006). The formal average uncertainty in the computed $12 + \log O/H$ values is typically 0.09 dex for the SDSS DR3 sample and 0.03 for the HeBCD sample. Concerning HII regions in spirals, many of which are located in the inner parts of the galaxies, only 22 objects out of over 300 have the intensity of $[O III] \lambda 4363$ measured. In some cases, other weak auroral lines have been measured. We have however decided not to adopt any T_e -based method, but rather to use the strong line method of Pilyugin (2001a) for high metallicities, which uses a combination of O_{23} and $[O III] \lambda 5007/[O II] \lambda 3727$ line ratios (the P -method). This method is based on a thorough consideration of many factors, it has been widely discussed (Pilyugin 2001b, 2003; Cédrés et al. 2004) and it is largely used. It is not clear at present whether more recent calibrations of the method (Pilyugin & Thuan 2005; Pilyugin et al. 2006) are more accurate, because of the problem of abundance bias at high metallicity discussed in Stasińska (2005). Since, in any event, abundances at high metallicities will have to be reconsidered once systematic measurements of mid-infrared lines become available, we simply took

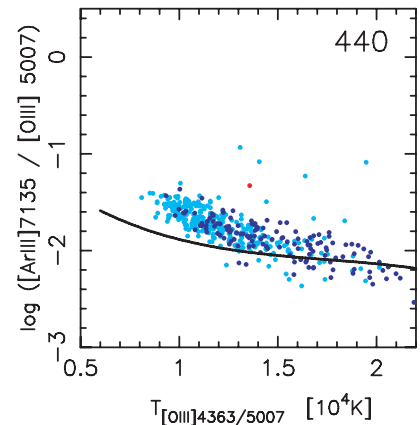


Fig. 2. $\log A_3O_3$ versus the electron temperature measured from the $[O III] \lambda 4363/5007$ ratio. Blue: HeBCD sample; cyan: SDSS DR3 sample. The curve shows the logarithm of the $[Ar III] \lambda 7135/[O III] \lambda 5007$ emissivity ratio at the temperature given by the abscissa (the curve has been shifted for easier comparison with observational points).

the most reasonable of the widely used methods, i.e. the calibration by Pilyugin (2001a). We discarded the objects for which the method gives $12 + \log O/H \leq 8.2$, which is the limit of validity defined by Pilyugin for the high metallicity range. Pilyugin considers his method to be accurate within 0.1 dex. However, the accuracy is certainly not as good near the bend in the O_{23} versus $12 + \log O/H$ relation.

3. Ar_3O_3 and S_3O_3 as metallicity calibrators

We show in Fig. 1a the values of $12 + \log O/H$ as a function of $\log Ar_3O_3$. Red dots correspond to HII regions from our spiral galaxies sample, blue ones correspond to the HeBCD sample, and cyan ones to the SDSS DR3 sample. In total, there are 666 objects in the plot. Apart from a few outliers in the SDSS DR3 sample (for which the oxygen abundance has likely been grossly underestimated, and which deserve detailed study), all the objects from the three samples merge into one monotonic sequence, as expected. The Spearman rank correlation coefficient, R_s , is 0.81. The observed trend of Ar_3O_3 is for a large part due to the decrease of T_e as metallicity increases, which translates into an increase of the $[Ar III] \lambda 7135$ and $[O III] \lambda 5007$ emissivity ratio (see Fig. 2). The fact that the hardness of the

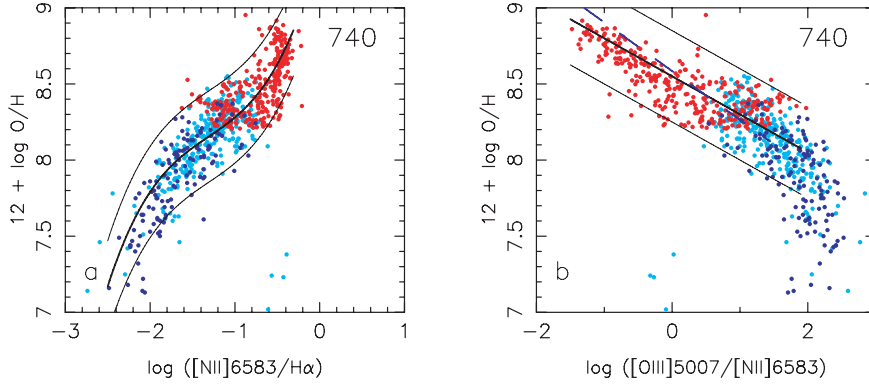


Fig. 3. **a)** $12 + \log O/H$ vs. $\log N_2$ for our data sample. The thick curve represents Eq. (3). **b)** $12 + \log O/H$ vs. $\log O_3N_2$. The dashed line shows the PP04 calibration. The thick line is our proposed calibration (Eq. (4)). The thin curves are vertically displaced from the thick one by ± 0.3 dex.

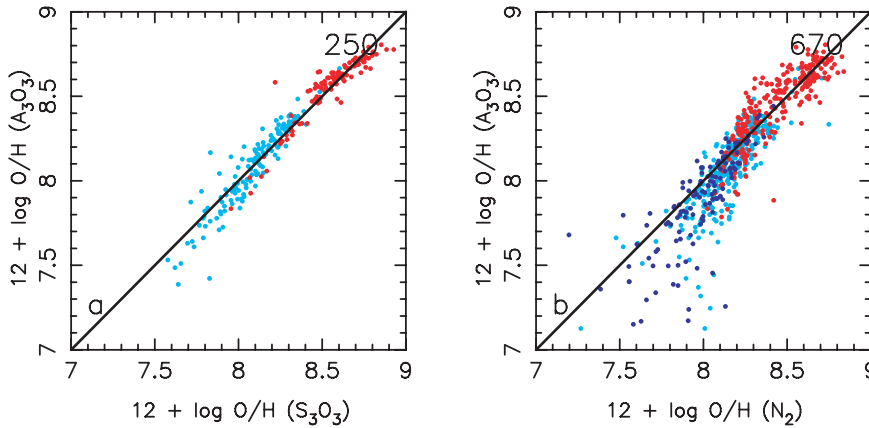


Fig. 4. **a)** $12 + \log O/H$ derived from Ar_3O_3 (Eq. (1)) vs. $12 + \log O/H$ derived from S_3O_3 (Eq. (2)). **b)** $12 + \log O/H$ derived from Ar_3O_3 (Eq. (1)) vs. $12 + \log O/H$ derived from N_2 using the formula of PP04 (Eq. (3)). The black lines show the 1:1 relations.

radiation field and ionization parameter decrease with increasing metallicity adds to the observed relation between $12 + \log O/H$ and $\log Ar_3O_3$, but does not play a dominant role. Also, as expected from stellar nucleosynthesis and confirmed by observations of metal-poor HII regions (Izotov et al. 2006), Ar and O are produced by the same stars, so, unlike O_3N_2 or N_2 , the Ar_3O_3 index is not affected by chemical evolution effects. Note that from Fig. 2 one can infer that the largest observed values of Ar_3O_3 seen in Fig. 1a must correspond either to a rapid change in the hardness of the radiation field and/or ionization parameter at the high metallicity end, or to a strong temperature gradient inside the HII regions.

From Fig. 1a, we can define a new metallicity calibration. The observed distribution of points in this figure is well fitted by the following expression:

$$12 + \log O/H = 8.91 + 0.34x + 0.27x^2 + 0.20x^3 \quad (1)$$

where $x = \log Ar_3O_3$. The heavy line in Fig. 1a represents this curve. The standard deviation between $\log O/H (O_{23})$ and $\log O/H (Ar_3O_3)$ using this curve is 0.23 dex. Of course, Eq. (1) may need revision when the physics of metal-rich HII regions becomes better understood.

Figure 1b shows the values of $12 + \log O/H$ as a function of $\log S_3O_3$. Here, the number of data points is much lower (283). Especially, there are no data for the HeBCD sample. But the Spearman rank correlation coefficient is still very high ($R_S = 0.84$). The trend shown by the observational points is similar to the one seen in Fig. 1a, with a smaller slope, due to the larger dependence of the ratio of [S III] $\lambda 9069$ and [O III] $\lambda 5007$

emissivities on electron temperature. The observed distribution of points can be fitted by:

$$12 + \log O/H = 8.70 + 0.28x + 0.03x^2 + 0.1x^3 \quad (2)$$

where $x = \log S_3O_3$. This is the heavy line in Fig. 1b. The standard deviation between $\log O/H (O_{23})$ and $\log O/H (S_3O_3)$ using this curve is 0.25 dex.

Both Eqs. (1) and (2) were found by eye estimate. We checked that, at least with our sample, the derivations of metallicity using both methods are consistent: there is no specific trend in the differences between both values of $\log O/H$ and the standard deviation of the differences is 0.09 dex (see Fig. 4a).

4. Discussion

We now discuss the relative merits of our Ar_3O_3 and S_3O_3 indexes with respect to O_3N_2 and N_2 . Rather than comparing our Figs. 1a and b with Figs. 1 and 2 of PP04 (who use a different data base and different derivations of O/H), we plot in Figs. 3a and b the values of $12 + \log O/H$ vs. $\log N_2$ and $\log O_3N_2$, respectively, using our samples. Note that our data points are more numerous (740 points instead of 137) and better distributed in the entire range of metallicities than the data points in PP04. The heavy line in Fig. 3a represents the calibration of the N_2 index as given by PP04:

$$12 + \log O/H = 9.37 + 2.03x + 1.26x^2 + 0.32x^3 \quad (3)$$

where $x = \log N_2$. As can be seen, their calibration fits our data well. On the other hand, their expression relating $12 + \log O/H$

with the O_3N_2 index, shown with the dashed blue line in Fig. 3b, is a very bad fit to our data. This stems from the fact that their relation is heavily weighted by only 4 objects in which the oxygen abundance has been determined from photoionization modelling and was found to be high. Unfortunately, abundances derived from models are very uncertain if constraints are insufficient. It is also very puzzling that those 4 points are quite separated from the rest of the data in the $(12 + \log O/H \text{ vs. } O_3N_2)$ diagram of PP04, while they are not in their $(12 + \log O/H \text{ vs. } N_2)$ diagram. Anyway, our Fig. 3b clearly confirms their conclusion that O_3N_2 alone is of no use for low metallicity objects. It can be used for high metallicity objects, though, rather than the calibration proposed by PP04 (their Eq. (3))² we prefer:

$$12 + \log O/H = 8.55 - 0.25x \quad (4)$$

where $x = \log O_3N_2$. However, comparison of Fig. 3b with Figs. 1a, b shows that there is no advantage in using O_3N_2 . Given the other shortcomings of O_3N_2 (see Sect. 1), it is preferable to use our new indexes, especially Ar_3O_3 , which does not require a large wavelength range. Note that [Ar III] was not considered as a “strong line” in previous studies but its intensity is comparable to that of [N II] at low metallicities. The Ar_3O_3 index is particularly well-suited to study the metallicities of the galaxies in the SDSS and at high redshifts.

The accuracy of the Ar_3O_3 , S_3O_3 and N_2 metallicity indicators is comparable at low metallicity, and the former two are more accurate at high metallicity, as can be judged from the distribution of points between the thin curves in Figs. 1a, b and 2a. But, as mentioned above, for *integrated galaxy spectra*, Ar_3O_3 and S_3O_3 should be superior even at low metallicity (at least they do not call for a different calibration to take into account the diffuse ionized medium). On the other hand, both Ar_3O_3 and S_3O_3 demand a reliable reddening correction, which is not required for N_2 . In the case of global spectra of galaxies, reddening correction can be done empirically by using the observed relation between the Balmer extinction and the galaxy excitation class or the discontinuity at 4000 Å (Stasińska et al. 2004). This avoids the measurements of $H\alpha$ and $H\beta$ I intensities, often somewhat problematic.

For the high metallicity end, our calibrations of Ar_3O_3 , S_3O_3 rely on the *P*-method of Pilyugin (2001a). When metal-rich HII regions become better understood, a recalibration of all the metallicity indexes in a consistent way may be needed. However, our discussion of the merits and drawbacks of Ar_3O_3 , S_3O_3 with respect to other methods will remain valid.

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² Note that PP04 use $([O III] \lambda 5007/H\beta)/([N II] \lambda 6584/H\alpha)$ while we defined O_3N_2 as $[O III] \lambda 5007/[N II] \lambda 6584$.