



HAL
open science

Investigating the nature and properties of MAXI J1810–222 with radio and X-ray observations

T.D. Russell, M. del Santo, A. Marino, A. Segreto, S.E. Motta, A. Bahramian,
S. Corbel, A. d’Aì, T. Di Salvo, J.C.A. Miller-Jones, et al.

► **To cite this version:**

T.D. Russell, M. del Santo, A. Marino, A. Segreto, S.E. Motta, et al.. Investigating the nature and properties of MAXI J1810–222 with radio and X-ray observations. *Monthly Notices of the Royal Astronomical Society*, 2022, 513 (4), pp.6196-6209. 10.1093/mnras/stac1332 . hal-03677050

HAL Id: hal-03677050








<https://hal.science/hal-03677050>

Submitted on 23 Mar 2023

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L’archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d’enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Investigating the nature and properties of MAXI J1810–222 with radio and X-ray observations

T. D. Russell ¹★, M. Del Santo ¹, A. Marino,^{1,2,3,4} A. Segreto,¹ S. E. Motta ⁵, A. Bahramian ⁶, S. Corbel,^{7,8} A. D’Aì,¹ T. Di Salvo ¹, J. C. A. Miller-Jones ⁶, C. Pinto,¹ F. Pintore ¹ and A. Tzioumis⁹

¹INAF, Istituto di Astrofisica Spaziale e Fisica Cosmica, Via U. La Malfa 153, I-90146 Palermo, Italy

²Dipartimento di Fisica e Chimica, Università degli Studi di Palermo, Via Archirafi 36, I-90123 Palermo, Italy

³Institute of Space Sciences (ICE, CSIC), Campus UAB, Carrer de Can Magrans S/N, E-08193 Barcelona, Spain

⁴Institut d’Estudis Espacials de Catalunya (IEEC), E-08034 Barcelona, Spain

⁵INAF, Osservatorio Astronomico di Brera, Via E. Bianchi 46, I-23807 Merate (LC), Italy

⁶International Centre for Radio Astronomy Research, Curtin University, GPO Box U1987, Perth, Western Australia 6845, Australia

⁷AIM, CEA, CNRS, Université de Paris, Université Paris-Saclay, F-91191 Gif-sur-Yvette, France

⁸Station de Radioastronomie de Nançay, Observatoire de Paris, CNRS, PSL Research University, Université d’Orléans, F-18330 Nançay, France

⁹Australia Telescope National Facility, CSIRO, P.O. Box 76, Epping, New South Wales 1710, Australia

Accepted 2022 May 9. Received 2022 May 8; in original form 2022 February 16

ABSTRACT

We present results from radio and X-ray observations of the X-ray transient MAXI J1810–222. The nature of the accretor in this source has not been identified. In this paper, we show results from a quasi-simultaneous radio and X-ray monitoring campaign taken with the Australia Telescope Compact Array, the Neil Gehrels *Swift* Observatory X-ray Telescope (*XRT*), and the *Swift* Burst Alert Telescope. We also analyse the X-ray temporal behaviour using observations from the Neutron star Interior Composition Explorer. Results show a seemingly peculiar X-ray spectral evolution of MAXI J1810–222 during this outburst, where the source was initially only detected in the soft X-ray band for the early part of the outburst. Then, ~ 200 d after MAXI J1810–222 was first detected the hard X-ray emission increased and the source transitioned to a long-lived (~ 1.5 yr) bright, harder X-ray state. After this hard state, MAXI J1810–222 returned back to a softer state, before fading and transitioning again to a harder state and then appearing to follow a more typical outburst decay. From the X-ray spectral and timing properties, and the source’s radio behaviour, we argue that the results from this study are most consistent with MAXI J1810–222 being a relatively distant ($\gtrsim 6$ kpc) black hole X-ray binary. A sufficiently large distance to source can simply explain the seemingly odd outburst evolution that was observed, where only the brightest portion of the outburst was detectable by the all-sky *XRT*s.

Key words: accretion, accretion discs – black hole physics – radio continuum: transients – X-rays: individual: MAXI J1810–222 – stars: neutron – X-rays: binaries.

1 INTRODUCTION

X-ray binaries (XRBs) consist of a compact object, either a stellar-mass black hole (BH) or a neutron star (NS), accreting matter from a companion star. Accretion typically occurs via Roche lobe overflow, where the infalling, accreted matter forms a differentially rotating accretion disc as it spirals in towards the compact object (e.g. Pringle, Rees & Pacholczyk 1973). However, not all of the infalling material is accreted on to the BH or NS primary star, instead some of it may be ejected from the system via outflows – either as a relativistic jet (e.g. Fender 2006) or disc winds (e.g. Díaz Trigo & Boirin 2016). While the accretion and wind phenomena are best observed in the optical and X-ray bands, the jets are observable at radio to infrared frequencies (and even up to X-rays and even gamma-rays).

Transient XRBs spend the majority of their lifetimes in a low-luminosity quiescent state (with an X-ray luminosity, L_X , of $\lesssim 10^{33}$

erg s^{−1}); however, they may occasionally go into phases of outburst due to increased mass accretion on to the compact object (thought to possibly arise from disc instabilities; e.g. van Paradijs & Verbunt 1984; Lasota 2001) that can last weeks to years (e.g. Tetarenko et al. 2016). During such outbursts their observable emission increases dramatically, where both the disc and jet emission increase by up to several orders of magnitude. Both BH and NS XRBs exhibit a variety of X-ray spectral states, broadly categorized as either hard or soft X-ray states. However, the outburst evolution for each source (or outburst) may vary (e.g. Done, Gierliński & Kubota 2007).

For BH XRBs, during a typical outburst the source is initially in a hard X-ray state, where the X-ray emission is dominated by a hard power law ($\Gamma < 2$) with a high-energy cutoff at ~ 50 – 100 keV. This component is usually interpreted as thermal Comptonization of soft disc photons by a population of hot electrons (*corona*, ~ 100 keV) (Zdziarski & Gierliński 2004). The hard state is associated with a flat to slightly inverted radio spectrum (with a radio spectral index $\alpha \gtrsim 0$, where the radio flux density, S_ν , is proportional to the observing frequency, ν , such that $S_\nu \propto \nu^\alpha$; e.g. Blandford & Königl 1979),

* E-mail: thomas.russell@inaf.it

arising from optically thick synchrotron emission originating from a persistent compact jet (e.g. Fender 2001, 2006; Corbel & Fender 2002).

As the outburst progresses, the source brightens as the mass accretion rate increases, but remains in the hard X-ray state. At the same time the radio emission also brightens. At some point, the X-ray emission begins to soften as a thermal component that can be modelled with a multitemperature blackbody (with an inner disc temperature, kT_{in} , of >0.3 keV), starts to become increasingly important. This emission is thought to be due to an optically thick, geometrically thin, accretion disc (Shakura & Sunyaev 1973).

As the disc luminosity (and temperature) increases, the X-ray emission softens. During this evolution the source can undergo transition through the hard and soft intermediate states (HIMS and SIMS, respectively) on its way to the thermal dominant soft state (e.g. Del Santo et al. 2008). This progression is accompanied by changes in the X-ray spectral and timing properties (e.g. Belloni & Motta 2016), as well as dramatic changes in the jet emission; during this progression, the compact jet switches off (Corbel et al. 2004; Fender, Belloni & Gallo 2004a; Fender 2010; Russell et al. 2020), being quenched by at least 3.5 orders of magnitude (Coriat et al. 2011; Russell et al. 2019; Carotenuto et al. 2021a) and a short-lived transient jet can be launched. The transient jet emission exhibits a steep radio spectrum ($\alpha \approx -0.6$) arising from ejected knots of (optically thin) synchrotron emitting plasma (e.g. Fender 2001), possibly as they collide with the pre-existing jet or the surrounding environment (e.g. Jamil, Fender & Kaiser 2010; Rushton et al. 2017), although that is still debated. These discrete ejecta are launched close in time to the transition from the hard state to the soft state (the precise connection to the changing accretion flow is not well understood; e.g. Fender, Homan & Belloni 2009; Miller-Jones et al. 2012; Tetarenko et al. 2017; Russell et al. 2019; Bright et al. 2020; Wood et al. 2021). Remaining in the soft state, the source then typically begins to fade as the mass accretion rate decreases. During this state, jet emission is generally not detected although the discrete ejecta may be detected away from the object as they travel outwards (e.g. Corbel et al. 2002; Russell et al. 2019; Bright et al. 2020; Carotenuto et al. 2021a).

As the source fades, the X-ray spectrum begins to harden once again, transitioning back to the hard state via the intermediate states in a reverse transition. Following its transition back to the hard state, the compact jet is progressively re-established over a period of several weeks (e.g. Miller-Jones et al. 2012; Corbel et al. 2013b; Kalemci et al. 2013; Russell et al. 2014). The outburst ends when the source fades further, returning to its quiescent state, showing a fading hard X-ray spectrum as it does so.

Transient NS XRBs generally show a broadly similar pattern of behaviour during their outbursts (e.g. Migliari & Fender 2006; Muñoz-Darias et al. 2014). However, the accretion–ejection picture is not much clear, possibly complicated by the presence of a stellar surface, boundary layer, and magnetosphere. The hard state for NS XRBs is associated with the launching of a compact jet, and these systems are capable of launching a transient jet during the hard-to-soft state transition (e.g. Fomalont, Geldzahler & Bradshaw 2001; Fender et al. 2004b; Migliari & Fender 2006). However, while the compact jet has been observed to quench in the soft state of some systems (Migliari et al. 2003; Tudose et al. 2009; Miller-Jones et al. 2010; Gusinskaia et al. 2017, 2020), it has not been identified in others (Rutledge et al. 1998; Kuulkers et al. 2003; Migliari et al. 2004; Migliari 2011). The cause for this discrepancy is unclear.

In their hard states, the radio (L_{R}) and X-ray (L_{X}) luminosities observed from BH XRBs appear to be coupled, showing a non-linear,

empirical correlation between the luminosities (e.g. Hannikainen et al. 1998; Corbel et al. 2000, 2003, 2013a; Gallo, Fender & Pooley 2003; Gallo, Miller & Fender 2012; Gallo, Degenaar & van den Eijnden 2018). For NS systems, the $L_{\text{R}}/L_{\text{X}}$ coupling is less straightforward, where individual systems can show various correlations, and not all sources follow a common or even defined track (e.g. Tudor et al. 2017; van den Eijnden et al. 2021). These variations may be a result of having an NS primary (with a surface and magnetic field) impacting the observed emission from the accretion flow (e.g. Maccarone 2008). It has also been suggested that NS XRB jets may show a different geometry or coupling to the accretion flow than in BH systems (Marino et al. 2020). NS XRBs are generally observed to be more radio faint than BH systems at a given X-ray luminosity (Fender & Kuulkers 2001; Migliari & Fender 2006), with the population typically being more radio faint by an average of ~ 22 (Gallo et al. 2018). Such a difference has been attributed to a number of different factors, such as jet power, primary mass, spin, magnetic field, and jet launching mechanism, although no direct connection has been identified yet (e.g. van den Eijnden et al. 2021). As such, although limited, a system’s radio and X-ray luminosities have often been used to identify the nature of the accretor in XRBs. Although, we note that some NS systems can be as radio bright as their BH counterparts at similar X-ray luminosities (e.g. Russell et al. 2018; Gusinskaia et al. 2020). NS XRBs can also be identified by the presence of thermonuclear Type-I X-ray bursts, which are sudden, short X-ray flashes that can be observed from accreting NS XRBs (Lewin, van Paradijs & Taam 1993). Type-I bursts are thought to arise from the ignition and burning of accreted matter that has accumulated on the NS surface (see Galloway & Keek 2021, for review).

1.1 MAXI J1810–222

MAXI J1810–222 was discovered on 2018 November 29 (MJD 58451; Negoro et al. 2018) by the Monitor of All-sky X-ray Image (MAXI) X-ray telescope on board the International Space Station (Matsuoka et al. 2009), being identified as a soft X-ray transient. *NuSTAR* observation was performed on MJD 58461 (2018 December 09), showing a soft X-ray spectrum with no apparent X-ray pulsations or bursts (Oeda et al. 2019). Due to Sun constraints, pointed X-ray observations did not occur again until the Neil Gehrels *Swift* Observatory X-ray Telescope (*XRT*) observed MAXI J1810–222 on 2019 February 09 (\sim MJD 58523). These first *Swift-XRT* observations measured the source position accurately and reported an X-ray spectrum that was still soft (Kennea & Negoro 2019). Analysis of ultraviolet (UV) and optical data from the *Swift Ultra-Violet and Optical Telescope* (*UVOT*) indicated a possible *Gaia* Data Release 2¹ counterpart, implying a source distance² of 730 ± 30 pc (Kennea & Negoro 2019), although this would result in a low-luminosity soft state.

Following the first *Swift-XRT* observation, MAXI J1810–222 faded steadily over the next ~ 30 d (until \sim MJD 58549). Early in this fading, when the source was still X-ray bright, radio observations taken on 2019 February 16 (MJD 58530) did not detect any radio emission from MAXI J1810–222 down to a 3σ upper limit of $99 \mu\text{Jy beam}^{-1}$ (Carotenuto et al. 2019). Around MJD 58700 (2019 August 05), analysis of *Burst Alert Telescope* (*BAT*) survey data with the *BAT-IMAGER* software (Segreto et al. 2010) showed an increase

¹Gaia Collaboration 2016, 2018.

²Bailer-Jones et al. 2018.

Table 1. Results from our ATCA radio observation. Each frequency band has a bandwidth of 2 GHz. Flux density errors include systematic uncertainties. Upper limits on the flux density are reported as $3 \times$ the RMS above the source position. We also provide the radio spectral index, α , defined as $S_\nu \propto \nu^\alpha$.

Date	MJD	Central frequency (GHz)	Flux density (μ Jy)	α
2019-11-13	58800.29 \pm 0.06	5.5	570 \pm 20	0.1 \pm 0.2
		9.0	576 \pm 16	
2020-12-13	59196.08 \pm 0.04	5.5	66 \pm 17	0.0 \pm 0.5
		9.0	65 \pm 15	
2020-12-27	59210.11 \pm 0.09	5.5	73 \pm 13	0.0 \pm 0.5
		9.0	73 \pm 12	
2021-01-05	59219.20 \pm 0.09	5.5	62 \pm 15	-0.2 \pm 0.5
		9.0	56 \pm 15	
2021-02-18	59263.80 \pm 0.06	5.5	70 \pm 20	0.3 \pm 0.5
		9.0	80 \pm 15	
2021-03-07	59280.03 \pm 0.09	5.5	75 \pm 15	-0.3 \pm 0.6
		9.0	60 \pm 20 ^a	
2021-04-18	59322.88 \pm 0.04	5.5	202 \pm 18	0.0 \pm 0.3
		9.0	200 \pm 16	
2021-04-30	59334.70 \pm 0.06	5.5	325 \pm 30	-0.07 \pm 0.30
		9.0	315 \pm 20	
2021-05-17	59351.90 \pm 0.03	5.5	240 \pm 30	0.14 \pm 0.30
		9.0	256 \pm 20	
2021-05-19	59353.86 \pm 0.04	5.5	265 \pm 40	0.4 \pm 0.3
		9.0	320 \pm 35	
2021-05-30	59365.84 \pm 0.06	5.5	120 \pm 30	1.3 \pm 0.7
		9.0	220 \pm 25	
2021-09-21	59478.31 \pm 0.05	5.5	208 \pm 16	-1.1 \pm 0.5
		9.0	125 \pm 18	
2021-10-16	59503.21 \pm 0.08	5.5	<75	-
		9.0	<70	
2021-11-13	59531.10 \pm 0.12	5.5	225 \pm 22	-0.4 \pm 0.3
		9.0	185 \pm 18	

Note.

^aThis 9-GHz observation suffered from some phase decorrelation due to poor weather, which we estimate to be of order 20 per cent and has been accounted for in the error (added in quadrature with the image and systematic errors).

in the hard X-ray emission from this source, and we triggered a radio monitoring campaign with the Australia Telescope Compact Array (ATCA). Over the next few years MAXI J1810–222 was reported to brighten and harden on MJD 58667 (2019 July 03; Negro et al. 2019), and again in 2020 late-February (from observations taken between MJDs 58906 and 58909; Ducci et al. 2020).

In this paper, we present radio and X-ray monitoring of this source taken over a span of more than 2 yr to identify the nature of the accretor. In Section 2, we describe the details of the radio and X-ray monitoring campaign. In Section 3, we outline the findings from these observations. In Section 4, we discuss the results, focusing on the nature of the compact object, its seemingly peculiar outburst evolution, and the distance to the source. We summarize our key findings in Section 5.

2 OBSERVATIONS AND DATA REDUCTION

2.1 ATCA radio observations

MAXI J1810–222 was observed by the ATCA 14 times between late-2019 and late-2021 (see Table 1). MAXI J1810–222 was first observed by ATCA on 2019 November 13 (MJD 58800) but not again until 2020 December, following which higher cadence radio

observations occurred between 2020 December and 2021 November. ATCA was in various configurations for these observations, but was mostly in its more extended 6 km, 1.5 km, and 750 m configurations.³ In all cases, the fixed location antenna 6 (located 6 km from the array core) was used during the analysis, providing angular resolutions of approximately a few arcseconds for all observations. Most observations were recorded simultaneously at central frequencies of 5.5 and 9 GHz, with a bandwidth of 2 GHz comprised of 2048 1-MHz channels at each band. However, observations taken on 2021 May 17 and 2021 May 19 were composed of 32 1-MHz channels spaced evenly (every 64 MHz) over 2 GHz of bandwidth as well as 2048 finer (31.25 kHz) channels within 64 MHz of bandwidth for both frequency bands.

For all observations, the flux and bandpass calibration was carried out using PKS 1934–638, while the nearby ($\sim 3.67^\circ$ away) J1817–254 was used for phase calibration. The data were edited for instrumental issues (shadowing, etc.) and radio frequency interference, calibrated, and imaged following standard procedures⁴ in the Common Astronomy Software Application (CASA version 5.1.2; McMullin et al. 2007). The calibrated data sets were imaged using the CASA task CLEAN. Imaging used a Briggs weighting scheme with a robust parameter of 0, balancing sensitivity and resolution.

The flux density, S_ν , of the source was measured using the CASA task IMFIT, where the flux density of the point source was determined by fitting an elliptical Gaussian with full width at half-maximum set by the synthesized beam shape. Errors on the absolute flux density scale include systematic uncertainties of 2 per cent for the 5.5-/9-GHz ATCA data⁵ (e.g. Massardi et al. 2011; McConnell et al. 2012), which were added in quadrature with the RMS of the image noise. The radio luminosity, L_R , was calculated by $L_R = 4\pi S_\nu \nu D^2$, where ν is the observing frequency and D is the distance to the source.

From our radio monitoring, we measure a best position (at 9 GHz) of the radio counterpart to MAXI J1810–222 at a right ascension (RA) and declination (Dec) of:

$$\text{RA (J2000)} = 18^{\text{h}} 12^{\text{m}} 39^{\text{s}}.76 \pm 0.11 \text{ arcsec}$$

$$\text{Dec (J2000)} = -22^\circ 19' 24''.92 \pm 0.26 \text{ arcsec},$$

where the errors are determined from the systematic errors (which are a function of the target’s distance from the phase calibrator) added in quadrature with the larger of either the theoretical error of beam centroiding (beam size/2 \times SNR, where SNR is the signal to noise ratio), which dominates the RA error, or the statistical error on the fit, which dominates the Dec error.

2.2 Swift

2.2.1 XRT

Swift-XRT carried out a monitoring campaign on MAXI J1810–222 starting from 2019 February 16 until November 2021, for a total of 41 visits (Table 2). The XRT observations were downloaded from the HEASARC public archive⁶ (source ID: 00011105) and processed with XRPIPELINE included within the HEASOFT software package (version 6.26.1). The latest version of CALDB available when these data were released⁷ was employed. We verified that the MAXI J1810–222 count rate in the observations performed in

³https://www.narrabri.atnf.csiro.au/operations/array_configurations/configurations.html

⁴https://casaguides.nrao.edu/index.php?title=Main_Page

⁵<https://www.atnf.csiro.au/observers/memos/d96783~1.pdf>

⁶<https://heasarc.gsfc.nasa.gov/docs/archive.html>

⁷caldb.indx20190910

Table 2. List of the *XRT* observations of the source used in this work. ObsIDs have been shortened for brevity, where * = 00011050. The MJD refers to the start of the X-ray observation.

ObsID	Date (UTC)	Date (MJD)	Obs. mode	Exposure (ks)
01	2019-02-09	58523.27	PC	0.96
02	2019-02-16	58530.25	WT	0.98
04	2019-02-19	58533.49	WT	0.99
05	2019-02-21	58535.95	WT	0.96
06	2019-02-23	58537.81	WT	0.99
07	2019-02-24	58539.00	WT	0.88
08	2019-03-03	58545.86	WT	0.96
09	2019-03-05	58547.57	WT	0.98
10	2019-03-07	58549.11	WT	1.06
11	2019-03-09	58551.49	WT	0.92
12	2019-03-11	58553.62	WT	0.38
13	2019-03-13	58555.35	WT	0.87
14	2019-07-05	58669.76	WT	0.99
15	2019-11-05	58792.80	PC	0.99
16	2021-02-09	59254.82	PC	1.00
17	2021-02-18	59263.84	PC	0.86
18	2021-02-25	59270.76	PC	0.93
19	2021-03-10	59283.63	PC	0.95
20	2021-03-28	59301.03	PC	0.89
21	2021-04-11	59315.71	WT	0.55
22	2021-04-15	59319.82	PC	0.99
23	2021-04-18	59322.39	PC	0.83
24	2021-04-21	59325.13	PC	1.00
25	2021-04-25	59329.96	PC	0.63
26	2021-04-25	59331.95	WT	0.42
27	2021-05-03	59337.08	WT	0.10
28	2021-05-09	59343.91	WT	0.97
29	2021-05-12	59346.51	PC	0.03
32	2021-05-23	59357.71	PC	0.09
33	2021-05-26	59360.50	PC	0.88
34	2021-05-30	59364.68	PC	0.92
36	2021-06-06	59371.65	PC	0.90
37	2021-08-20	59446.92	PC	0.18
40	2021-09-08	59465.96	PC	0.89
41	2021-09-16	59473.79	WT	0.20
43	2021-09-21	59478.32	PC	0.86
44	2021-09-30	59487.28	PC	0.97
45	2021-10-06	59493.31	PC	0.91
46	2021-10-13	59500.55	PC	1.01
47	2021-10-20	59507.58	PC	1.36
48	2021-10-27	59514.56	PC	0.87
49	2021-11-03	59521.32	PC	1.44

window timing (WT) mode was below 100 cts s⁻¹, thus resulting in a negligible pile-up impact. Observations in photon counting (PC) mode were moderately affected by pile-up ($\gtrsim 0.5$ cts s⁻¹). With DS9 we used circles of 20 pixels⁸ (corresponding to ~ 47 arcsec) for the source and background regions, while when pile-up was an issue (in PC mode) annuli were used where the inner and outer radius were dependent on the count rate (see <https://heasarc.gsfc.nasa.gov/lheasoft/ftools/headas/xrtgrblc.html> for discussion). These regions were centred on the measured *XRT* position of RA (J2000) = 18^h12^m39^s.66, Dec (J2000) = -22°19′25″.0 (with an error radius of 1.8 arcsec after *Swift-UVOT* enhancement, consistent with our measured radio position). We processed the light curves and spectra

by running the task XRTPRODUCTS. We rebinned each spectrum with GRPPHA in order to have at least 25 counts per bin, which allows the use of χ^2 statistics. Due to the relatively high column density and brightness of the source, we excluded data below 0.8 keV. To explore if short-time-scale X-ray flares and bursts were present in the observations we carefully checked the light curve of each observation, searching for intra-observational variability.

2.2.2 BAT

Data from the *BAT* survey were retrieved from the HEASARC public archive and processed using the *BAT-IMAGER* software (Segreto et al. 2010). This code processes data collected by coded mask telescopes. It computes all-sky maps after modelling and subtracting the proper background. Then, for each detected source, light curves and spectra can be extracted.

We produced light curves of MAXI J1810–222 in count rate per pixel in the energy range 15–50 keV, with 15-d binning time (Fig. 1), as well as the 50- to 80-keV range. Spectra were extracted in the range 15–90 keV, with logarithmic binning (for a total of six bins) and the official *BAT* spectral redistribution matrix was used. We checked to ensure that no other sources within the *BAT* field of view contaminated our results.

2.3 NuSTAR

NuSTAR observed MAXI J1810–222 on MJD 58461 (2018 December 9), with two pointings (ObsID 90410366001 and 90410351001) of exposures of 0.8 and 12.7 ks, respectively. Among them, we analysed only the observation with higher exposure. Data were reduced using the standard NUSTARDAS task, incorporated in HEASOFT (version 6.26.1). The observation was taken in observing mode ‘06’, which corresponds to the case when the on-board star tracker, the Camera Head Unit (CHU) #4, is not available because it was either blinded by a bright source or by the Earth itself, so that the image reconstruction accuracy of *NuSTAR* becomes profoundly degraded. The image of the source displays multiple centroids, so that an additional procedure is necessary to obtain reliable data. We used the task NUSPLITSC on cleaned event data in order to split the event files in several event files for each CHU combination (CHUs 3, 1+2, and 2 + 3 were used). After the treatment, we checked that the profile of the source was well defined and that it displayed a single centroid. For each event file, we extracted scientific products with the command NUPRODUCTS by using a circular region with a 60 arcsec radius as the source region (centred at the radio position of the source), sufficient to encapsulate the source point spread function (PSF). In order to take into account any background non-uniformity on the detector, for the background region we used four circular areas of 50 arcsec in radius, placed on areas of the image with negligible contamination from the source. Finally, we summed the spectra corresponding to each CHU combination in a single spectrum. We repeated the whole procedure for both the two hard X-ray imaging telescopes on board *NuSTAR*, i.e. the focal plane mirror (FPM) A and B. The last step was to re-bin the spectra, in order to have at least 25 counts per bin.

2.4 Neutron star Interior Composition Explorer

Neutron star Interior Composition Explorer (NICER) observed MAXI J1810–222 84 times between February 2019 and November 2020. A complete NICER spectral and timing analysis of MAXI J1810–222 is beyond the scope of this work (and will be

⁸As suggested by the *Swift-XRT* guidelines, e.g. <https://www.swift.ac.uk/analysis/xrt/spectra.php>.

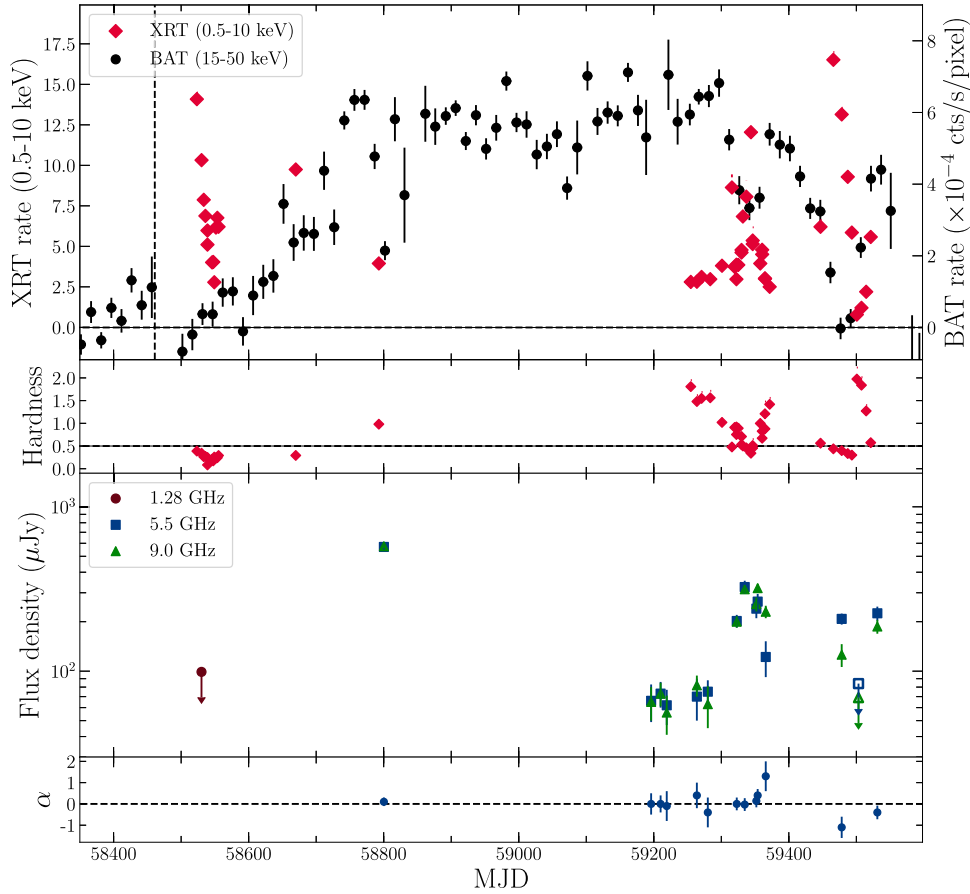


Figure 1. X-ray and radio light curves of MAXI J1810–222. First panel: *Swift-XRT* (red diamonds) and *BAT* (black circles) X-ray light curves. The horizontal dashed line marks zero and the vertical dashed line represents the timing of the *NuSTAR* observation. Second panel: X-ray hardness from the *Swift-XRT* monitoring, where hardness is defined as $(2\text{--}10\text{ keV})/(0.5\text{--}2\text{ keV})$. As a guide, we arbitrarily place a horizontal line at a hardness of 0.5. Third panel: Multifrequency radio light curve, where the 5.5- and 9-GHz data are shown as the blue squares and green triangles, respectively. The brown circle also shows the soft-state 1.28-GHz non-detection reported by Carotenuto et al. (2019), where we show the 3σ upper limit. Fourth panel: Radio spectral index, α . The long-term X-ray and radio monitoring shows a seemingly atypical evolution during MAXI J1810–222’s long-lived outburst.

presented in a forthcoming paper). As such, here we only selected three observations representative of the different spectral states we observed, in order to explore the overall variability properties of the source: We studied NICER observations with the identification numbers 1200560105, 2200560121, and 2200560140. These observations were selected to be as close as possible to the *XRT* spectra with ObsIDs 0001105002, 0001105014, and 0001105015 (see Section 3).

For each of the above observations we visually inspected the light curves and determined appropriate good time intervals (GTIs) to remove instrumental artefacts. Then we extracted power density spectra (PDS) in the 0.5- to 15.0-keV energy range in segments ≈ 15 -s long. We then averaged the Leahy-normalized PDS created from each segment to produce one averaged PDS per observation with a Nyquist frequency of ≈ 1000 Hz. We did not correct for the contribution of the Poisson noise, but fitted it with a flat power law when modelling the PDS.

3 RESULTS

In Fig. 1, we show the X-ray and radio light curves, X-ray hardness (defined as the count rate ratio 2–10 keV/0.5–2 keV) and radio spectral index, α , obtained from our long-term monitoring

Table 3. List of the *BAT* observations of the source used in this work. The MJD ranges represent the time ranges from which survey *BAT* spectra were extracted. We also report the corresponding *XRT* ObservationsID performed within the *BAT* time range, where the *XRT* ObsIDs are given with the prefix of 00011050 removed (* = 00011050).

Start date		<i>BAT</i> End date		Exp. (ks)	<i>XRT</i> ObsID
(UTC)	(MJD)	(UTC)	(MJD)		
2019-06-06	58640.0	2019-08-03	58699.0	537	*14
2019-10-04	58760.0	2020-01-01	58849.0	718	*15
2021-01-10	59224.0	2021-03-10	59284.0	745	*17
2021-03-27	59300.0	2021-04-26	59330.0	365	*21
2021-04-26	59330.0	2021-05-16	59350.0	250	*28
2021-05-30	59364.0	2021-07-13	59408.0	469	*34

campaign. The spectral evolution of the system is also shown by the *XRT* hardness intensity diagram (HID; Fig. 2). This HID has been computed by estimating the 0.5- to 2-keV/2- to 10-keV fluxes with a spectral fitting procedure. Using *xspec*, we modelled each *XRT* spectrum with either a multicolour disc blackbody (*discbb*) or a thermal Comptonization model (*nthcomp*), or a

Table 4. Fit results of the *Swift*/*XRT* spectra. We used `discbb`, `nthcomp`, or both when required by the data. The `tbabs` model has been applied to all the spectra to take into account the neutral absorption. The electron temperature of the Comptonization component, kT_e , has been fixed at 25 keV. Parameters in round parentheses were kept frozen. Quoted errors reflect 90-per-cent confidence level. For the ObsID * = 00011050.

Obs.	N_{H} ($\times 10^{22} \text{ cm}^{-2}$)	kT_{disc} (keV)	K_{disc}	Γ	$F_{0.5-10 \text{ keV}}$ ($\times 10^{-10} \text{ erg cm}^{-2} \text{ s}^{-1}$)	χ^2_{ν} (degree of freedom)
01	$1.34^{+0.14}_{-0.12}$	0.53 ± 0.03	770^{+300}_{-210}	–	$3.44^{+0.14}_{-0.11}$	0.96(76)
*02	1.38 ± 0.07	$0.493^{+0.014}_{-0.013}$	830^{+160}_{-130}	–	2.45 ± 0.05	1.13(177)
*04	$1.24^{+0.09}_{-0.08}$	0.45 ± 0.02	980^{+240}_{-190}	–	1.93 ± 0.04	0.76(150)
*05	1.21 ± 0.08	$0.45^{+0.02}_{-0.01}$	840^{+210}_{-160}	–	1.56 ± 0.04	1.01(137)
*06	1.40 ± 0.20	$0.42^{+0.04}_{-0.03}$	630^{+480}_{-260}	–	$0.73^{+0.06}_{-0.04}$	0.91(36)
*07	$1.27^{+0.10}_{-0.09}$	0.43 ± 0.02	930^{+280}_{-210}	–	1.37 ± 0.04	1.01(124)
*08	$1.27^{+0.13}_{-0.12}$	0.41 ± 0.02	790^{+390}_{-240}	–	$0.89^{+0.04}_{-0.03}$	1.06(93)
*09	$1.14^{+0.12}_{-0.11}$	0.40 ± 0.02	870^{+370}_{-250}	–	$0.87^{+0.04}_{-0.03}$	0.98(99)
*10	$1.46^{+0.16}_{-0.15}$	0.41 ± 0.02	880^{+460}_{-290}	–	$0.86^{+0.05}_{-0.04}$	0.94(72)
*11	1.15 ± 0.09	0.45 ± 0.02	700^{+210}_{-160}	–	1.42 ± 0.04	0.96(131)
12	$1.22^{+0.15}_{-0.13}$	0.43 ± 0.02	1360^{+670}_{-420}	–	$2.04^{+0.08}_{-0.07}$	1.07(58)
13	$1.43^{+0.11}_{-0.10}$	0.45 ± 0.02	850^{+270}_{-200}	–	1.48 ± 0.05	0.87(127)
14	0.74 ± 0.05	0.38 ± 0.02	1800^{+920}_{-570}	$2.1^{+0.3}_{-0.1}$	$2.98^{+0.11}_{-0.10}$	0.91(200)
*15	0.4 ± 0.01	–	–	1.6 ± 0.1	$2.10^{+0.20}_{-0.15}$	1.17(55)
*16	1.0 ± 0.3	–	–	$1.38^{+0.19}_{-0.17}$	2.50 ± 0.20	0.93(42)
*17	1.0 ± 0.2	–	–	$1.63^{+0.07}_{-0.06}$	1.80 ± 0.20	0.77(45)
*18	1.3 ± 0.3	–	–	1.7 ± 0.2	$1.80^{+0.15}_{-0.17}$	0.95(40)
*19	$1.04^{+0.12}_{-0.10}$	–	–	1.47 ± 0.05	$1.87^{+0.19}_{-0.16}$	0.78(41)
*20	0.70 ± 0.20	–	–	1.58 ± 0.20	2.00 ± 0.20	1.10(38)
*21	$1.08^{+0.15}_{-0.12}$	0.32 ± 0.02	3825^{+4230}_{-2250}	$1.8^{+0.1}_{-0.2}$	$3.60^{+0.30}_{-0.20}$	0.92(131)
*22	0.9 ± 0.2	–	–	1.9 ± 0.2	$1.70^{+0.19}_{-0.16}$	0.78(28)
*23	$0.6^{+0.3}_{-0.2}$	–	–	1.6 ± 0.2	1.50 ± 0.20	0.81(21)
*24	1.0 ± 0.3	–	–	2.0 ± 0.2	$1.62^{+0.18}_{-0.15}$	0.78(28)
*25	1.0 ± 0.3	–	–	2.2 ± 0.3	1.80 ± 0.20	0.78(28)
*26	1.0 ± 0.2	$0.31^{+0.08}_{-0.07}$	2700^{+1200}_{-1900}	1.8 ± 0.4	2.70 ± 0.20	0.90(81)
*27	$0.90^{+0.40}_{-0.30}$	–	–	2.8 ± 0.5	$3.60^{+0.60}_{-0.40}$	0.90(14)
*28	1.13 ± 0.08	0.35 ± 0.03	3895^{+2330}_{-1330}	2.1 ± 0.3	$3.67^{+0.12}_{-0.11}$	1.02(211)
*29	(1.0)	$0.35^{+0.07}_{-0.06}$	1310^{+1600}_{-710}	<2.1	$2.50^{+0.50}_{-0.30}$	0.90(22)
*33	0.9 ± 0.2	–	–	2.2 ± 0.2	$1.62^{+0.18}_{-0.14}$	1.00(29)
*34	0.8 ± 0.2	–	–	1.8 ± 0.2	$1.49^{+0.18}_{-0.14}$	0.89(24)
*36	0.7 ± 0.2	–	–	$1.7^{+0.3}_{-0.2}$	$1.44^{+0.16}_{-0.18}$	0.98(20)
*40	$1.43^{+0.16}_{-0.15}$	0.54 ± 0.03	990^{+430}_{-290}	–	$4.65^{+0.18}_{-0.16}$	1.21(71)
*41	1.02 ± 0.11	0.55 ± 0.03	660^{+240}_{-170}	–	4.15 ± 0.16	0.99(72)
43	$1.41^{+0.19}_{-0.16}$	0.52 ± 0.04	1040^{+570}_{-350}	–	$3.77^{+0.18}_{-0.17}$	1.20(55)
*44	$1.31^{+0.16}_{-0.15}$	0.49 ± 0.03	920^{+440}_{-280}	–	2.59 ± 0.11	0.87(56)
*45	$1.35^{+0.20}_{-0.18}$	0.44 ± 0.03	900^{+660}_{-350}	–	1.42 ± 0.08	0.94(37)
*47	(1.40)	–	–	1.7 ± 0.2	$0.72^{+0.09}_{-0.08}$	0.78(13)
*48	1.20 ± 0.50	–	–	1.9 ± 0.3	$1.10^{+0.14}_{-0.15}$	0.95(15)
*49	$0.80^{+0.30}_{-0.20}$	$0.47^{+0.18}_{-0.15}$	170^{+870}_{-120}	$1.8^{+0.5}_{-0.8}$	$2.12^{+0.16}_{-0.15}$	1.3(57)

combination of the two models when required. In the spectral fits which included `nthcomp`, the `np_type` parameter was set to 1, i.e. corresponding to disc blackbody seed photons, and the seed photons temperature tied to the kT_{disc} parameter when `discbb` was included or fixed to 0.1 when no disc component was required. For all spectra, the interstellar absorption was left free and fitted with the model `tbabs`, using photoelectric cross-sections from Verner

et al. (1996) and element abundances from Wilms, Allen & McCray (2000). Spectral parameters of each *XRT* spectrum are shown in Table 4.

As apparent from Fig. 2, the system displayed an atypical spectral evolution along the outburst. Based on the spectral variability, we selected a number of *XRT* spectra and combined them with quasi-simultaneous *BAT* spectra (Table 3). In the following sections,

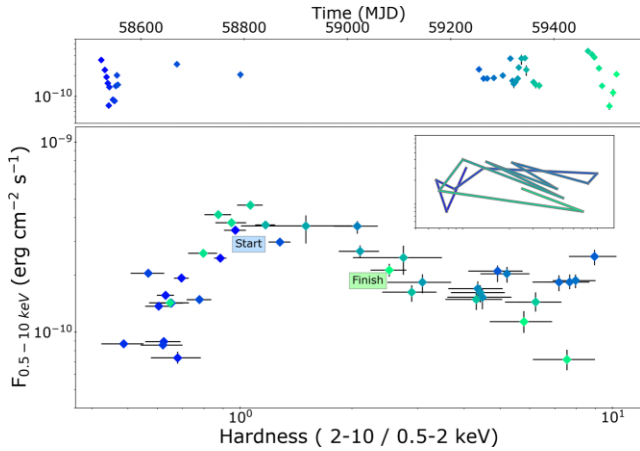


Figure 2. *XRT* light curve (top panel) and HID (lower panel) of MAXI J1810–222. The time sequencing of the data is highlighted with a colour map, from blue to green as shown in the top panel. The inset tracks out the hardness evolution during the outburst. Throughout its outburst, MAXI J1810–222 was observed to transition back and forth between bright softer and harder X-ray states, peculiar for a BH or NS XRB.

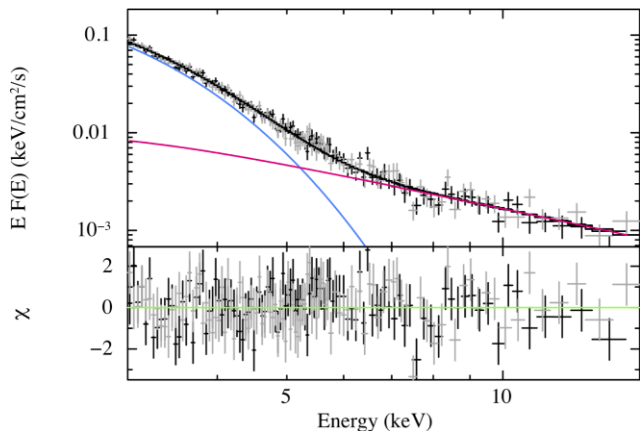


Figure 3. Residuals for the *NuSTAR* observation performed during the soft state of the system. Data: FPM A (black) and FPM B (grey). The blue (magenta) thick line indicates the blackbody (`nthcomp`) contributions to the total spectrum. The *NuSTAR* spectrum was well described by a disc blackbody with a (steep) hard X-ray power-law component.

we report on the broad-band spectral fitting results as well radio evolution during the outburst.

3.1 Outburst evolution in the X-ray and radio

3.1.1 Initial discovery in the soft state

The discovery X-ray detection and X-ray follow-up soon after its initial detection implied a soft X-ray spectrum, as reported by both *MAXI* and *NuSTAR* teams (Matsuoka et al. 2009; Oeda et al. 2019). Indeed, the *NuSTAR* spectrum (see Fig. 3) can be well described by a ~ 0.5 -keV disc blackbody and a steep hard X-ray tail with $\Gamma \sim 3.5$ (see Table 5 for details on the spectral fitting parameters). In addition, MAXI J1810–222 was not significantly detected with the *BAT* hard X-ray telescope (Fig. 1), confirming that the system was in a softer X-ray state. However, it is noteworthy that there is a hint of a potential brightening in stacked *BAT* observations from early

2018 November (from MJD 58426.6 ± 7.5) – in the weeks before the first reported X-ray detection of the source (see Fig. 1 showing the marginal rise).

The first *XRT* observation occurred on 2019 February 09 (MJD 58523), showing bright X-ray emission (6×10^{-10} erg s^{-1} in the 0.5- to 10-keV band; Figs 1–2). The spectrum of this first observation is well described with a simple `discbb` model, with a disc temperature consistent with the *NuSTAR* value (Fig. 4, top-left panel, and Table 5). The *NICER* observation (ObsID 1200560101), taken 2 d after the first *XRT* observation, shows an RMS consistent with zero (< 2 per cent), with a PDS showing essentially only Poisson noise (see Fig. 4, top-right panel).

The 0.5- to 10-keV emission then faded over the next few months (see dark blue points in Fig. 2). Throughout this decay the X-ray spectra were well described with a disc blackbody model ($kT_{\text{disc}} \sim 0.4$ – 0.5 keV) and a hardness ratio of < 0.5 (Fig. 1), implying the source being in soft state. Indeed, during this phase of the outburst, no radio counterpart to MAXI J1810–222 was detected during 1.28-GHz MeerKAT observations taken on MJD 58530, with a 3σ upper limit of $99 \mu\text{Jy beam}^{-1}$ (Carotenuto et al. 2019).

3.1.2 Brightening of the hard X-ray emission

Around MJD 58600, the 15- to 50-keV emission was first detected, brightening considerably over the next few months before levelling off around MJD 58750. The first *XRT* observation taken during this brightening in hard X-rays occurred on MJD 58670 (ObsID 00011105014). We extracted a *BAT* spectrum averaged over an interval of ~ 30 d centred at the time of this *XRT* observation, as no significant hard variability was observed in the *BAT* light curve during this time. We fitted these *XRT* and *BAT* spectra simultaneously. Significant emission is detected up to ~ 100 keV (Fig. 4, middle-left panel) and is well described by thermal Comptonized emission (`nthcomp`) plus a thermal, standard, disc emission (`discbb`). $\Gamma = 2.1$ and a lower limit of 30 keV for the electron temperature were obtained, as the disc temperature decreased, i.e. to $kT_{\text{disc}} \sim 0.4$ keV (see Table 5). A quasi-simultaneous *NICER* observation performed on MJD 58672 (ID 2200560121) showed an increased RMS of ≈ 14 per cent, and a weak flat-top noise PDS (Fig. 4, middle-right panel), consistent with an X-ray hardening of the source. No radio observations of MAXI J1810–222 were taken during this phase of the outburst.

Around MJD 58800 the hard X-ray emission was observed to briefly fade (for a few weeks). A single *XRT* observation showed that the soft emission had also faded around this time, when compared to the previous *XRT* observation. Despite the decreasing luminosity in both the hard and soft X-ray band, the broad-band X-ray spectrum was harder (Fig. 4, bottom-left panel), as shown by the increase in the X-ray hardness (Fig. 1). At this time, the disc emission becomes negligible (see Table 5), and the Comptonization component flattened ($\Gamma \sim 1.6$), suggesting that the source was in a hard X-ray state. In addition, *NICER* observations showed the fractional X-ray emission variability progressively increased during this period, as expected for a source transitioning from the soft state, towards and then into a hard X-ray spectral state. While no X-ray quasi-periodic oscillations (QPOs) were detected from MAXI J1810–222, *NICER* observations on MJD 58793 (ID 2200560140) showed an RMS of ≈ 25 per cent and a QPO-free flat-top noise PDS, breaking at ≈ 1 Hz (Fig. 4, bottom-right panel).

ATCA observations taken on MJD 58800 detected the relatively bright radio counterpart to MAXI J1810–222 (Table 1). We mea-

Table 5. Fitting results of different spectral states. Parameters in round parentheses were kept frozen. Quoted errors reflect 90-percent confidence level. Observations are labelled for ObsID, †=90402370, * = 00011050. Parameters without a component for that fit are marked by ‘–’.

Observatory	Observation	N_{H}	kT_{in}	K_{disc}	Γ	kT_{e}	$F_{\text{unabs, 0.1–100 keV}}$		χ^2_{ν}
		($\times 10^{22} \text{ cm}^{-2}$)	(keV)			(keV)	($\times 10^{-9} \text{ erg cm}^{-2} \text{ s}^{-1}$)		(degree of freedom)
							discbb	nthcomp	
<i>NuSTAR</i>	†02	(1.0)	0.49 ± 0.01	780^{+160}_{-130}	$3.45^{+0.2}_{-0.25}$	>10	0.958 ± 0.020	0.050 ± 0.005	1.04(217)
<i>XRT</i>	*01	$1.34^{+0.14}_{-0.12}$	0.53 ± 0.03	770^{+300}_{-210}	–	–	1.410 ± 0.120	–	0.96(76)
<i>XRT + BAT</i>	*14	0.74 ± 0.05	0.38 ± 0.02	1800^{+920}_{-570}	$2.1^{+0.3}_{-0.1}$	>30	0.710 ± 0.040	$0.320^{+0.060}_{-0.030}$	0.91(200)
<i>XRT + BAT</i>	*15	0.4 ± 0.1	–	–	1.6 ± 0.1	24^{+16}_{-5}	–	$0.960^{+0.180}_{-0.160}$	1.17(55)
<i>XRT + BAT</i>	*17	1.0 ± 0.2	–	–	$1.63^{+0.07}_{-0.06}$	>54	–	$0.80^{+0.10}_{-0.09}$	0.77(45)
<i>XRT + BAT</i>	*21	$1.08^{+0.15}_{-0.12}$	0.32 ± 0.02	3825^{+4230}_{-2250}	$1.8^{+0.1}_{-0.2}$	>30	$0.720^{+0.080}_{-0.070}$	$1.050^{+0.400}_{-0.200}$	0.92(131)
<i>XRT + BAT</i>	*28	1.13 ± 0.08	0.35 ± 0.03	3895^{+2330}_{-1330}	2.1 ± 0.3	(20)	1.170 ± 0.060	$0.650^{+0.040}_{-0.030}$	1.02(211)
<i>XRT + BAT</i>	*34	0.6 ± 0.1	–	–	$1.90^{+0.08}_{-1.0}$	>48	–	0.60 ± 0.10	0.71(37)
<i>XRT</i>	*41	1.02 ± 0.11	0.55 ± 0.03	660^{+240}_{-170}	–	–	$1.32^{+0.13}_{-0.12}$	–	0.99(72)

sured radio flux densities of $570 \pm 20 \mu\text{Jy}$ at 5.5 GHz and $576 \pm 16 \mu\text{Jy}$ at 9 GHz, providing a radio spectral index, α , of 0.1 ± 0.2 , consistent with a flat to mildly inverted radio spectrum (Fig. 1).

MAXI J1810–222 remained quite stable in the 15- to 50-keV X-rays band over the following ~ 1.5 yr. Towards the end of this phase, higher cadence *XRT* observations also show relatively stable (albeit faint) 0.5- to 10-keV emission, with hard X-ray spectra (X-ray hardness exceeding 1 and lack of the disc blackbody component, see Table 4) indicating that the system was still in the hard state (see e.g. Obs. *17 in Table 5).

We also extracted the *BAT* light curve in the 50- to 80-keV energy band (see Fig. 5). While the 50- to 80-keV X-rays typically traced a similar behaviour to the 15- to 50-keV X-rays, we did observe some bright flares in the 50- to 80-keV band that were not seen in the lower X-ray energies. In particular, during the time intervals $\sim \text{MJD } 58850\text{--}59000$ and $\sim \text{MJD } 59150\text{--}59300$ the hardest 50–80 keV showed significant brightening above the 15- to 50-keV X-rays. Such flares in only the hardest X-ray bands can arise from either an increase of the electron temperature of the thermal Comptonization, or from the appearance of an additional high energy component.

It is worth noting that the ATCA radio observations taken during the $\sim \text{MJD } 59150\text{--}59300$ range showed steady, fainter radio emission, with flux densities of $\sim 60\text{--}80 \mu\text{Jy}$. Over this time the radio spectrum remained consistent with flat ($\alpha \approx 0$), although due to the faintness of the radio counterpart the errors on the radio spectrum were large.

Around MJD 59322 the hard X-ray emission once again briefly faded, decreasing by a factor of ~ 2 over a period of a few weeks. This time, however, it was accompanied by a brightening of the soft 0.5- to 10-keV emission (which increased by a factor of ~ 4 ; Fig. 1). This X-ray softening is apparent in both the X-ray hardness (Fig. 1) and the HID (Fig. 2). In the first *XRT* observation (ObsID *21) taken during this phase, the contribution from a cold disc ($kT_{\text{disc}} \sim 0.3$ keV) becomes significant again and Γ steepened to ~ 1.8 . The softer X-ray spectral behaviour exhibited during this brightening was considerably variable (see Fig. 2), with the system erratically moving back and forth from high flux softer states (such as Obs. *28, Table 5) to low flux harder states (Obs. *34, Table 5). Only the higher X-ray flux observations require the addition of a cold disc in the X-ray spectral models (see Table 4). As the X-rays brightened, the radio emission also increased. Throughout the brightening, the radio spectrum remained flat ($\alpha \approx 0$). Interestingly, as the source once

again faded to its pre-brightening levels (from a few weeks earlier) the radio spectrum became increasingly inverted ($\alpha = 1.3 \pm 0.7$).

3.1.3 Hard X-ray decline

The hard X-rays began to fade away after $\sim \text{MJD } 59400$. Around the same time, the soft X-rays also brightened considerably, implying a return to a soft X-ray state (Figs 1 and 2). Around the time of the *XRT* peak (in Obs. *41) the hard X-ray emission (beyond 15 keV) became too faint to be detected. The *XRT* spectrum can be well described by a simple discbb model, with $kT_{\text{disc}} \sim 0.6$ keV, the highest disc temperature observed up to this point. The radio spectrum during the decline was measured to be steep ($\alpha = -1.1 \pm 0.5$).

The *XRT* emission then faded and the spectrum hardened. This coupled with the X-ray hardness and the source evolution in the HID suggests a return to the hard state. No radio emission was detected from MAXI J1810–222 during this hard state.

An additional ATCA observation taken a few weeks later showed relatively bright radio emission with a radio spectrum consistent with flat (Table 1 and Fig. 1). At the same time, *BAT* monitoring showed that the 15- to 50-keV emission had once again brightened suggesting a return to a bright hard state. However, that return was short-lived and the source faded, becoming undetected by the *BAT* monitoring suggesting a possible end to the outburst (at least during our monitoring). Although, due to the lack of further monitoring at soft X-rays, it is also possible that the source remained in a bright soft state beyond this point.

3.2 Radio position and reported possible *Gaia* counterpart

Our measured radio position (Section 3.2) is consistent with the *Swift-XRT* position (Kennea & Negoro 2019), as shown in Fig. 6. From optical and UV observations, Kennea & Negoro (2019) proposed a possible *Gaia* counterpart to MAXI J1810–222 (labelled as G2 in the lower panel of Fig. 6), although this possible counterpart was slightly outside the *XRT* 90-percent error region. If G2 is the optical counterpart to MAXI J1810–222, it would imply a relatively low source distance of 730 ± 30 pc. However, our radio source position is not consistent with the potential optical/UV counterpart, G2, or any other source listed in the *Gaia* catalogue, where the nearest object, labelled G1, in Fig. 6, lies ≈ 3 arcsec from our

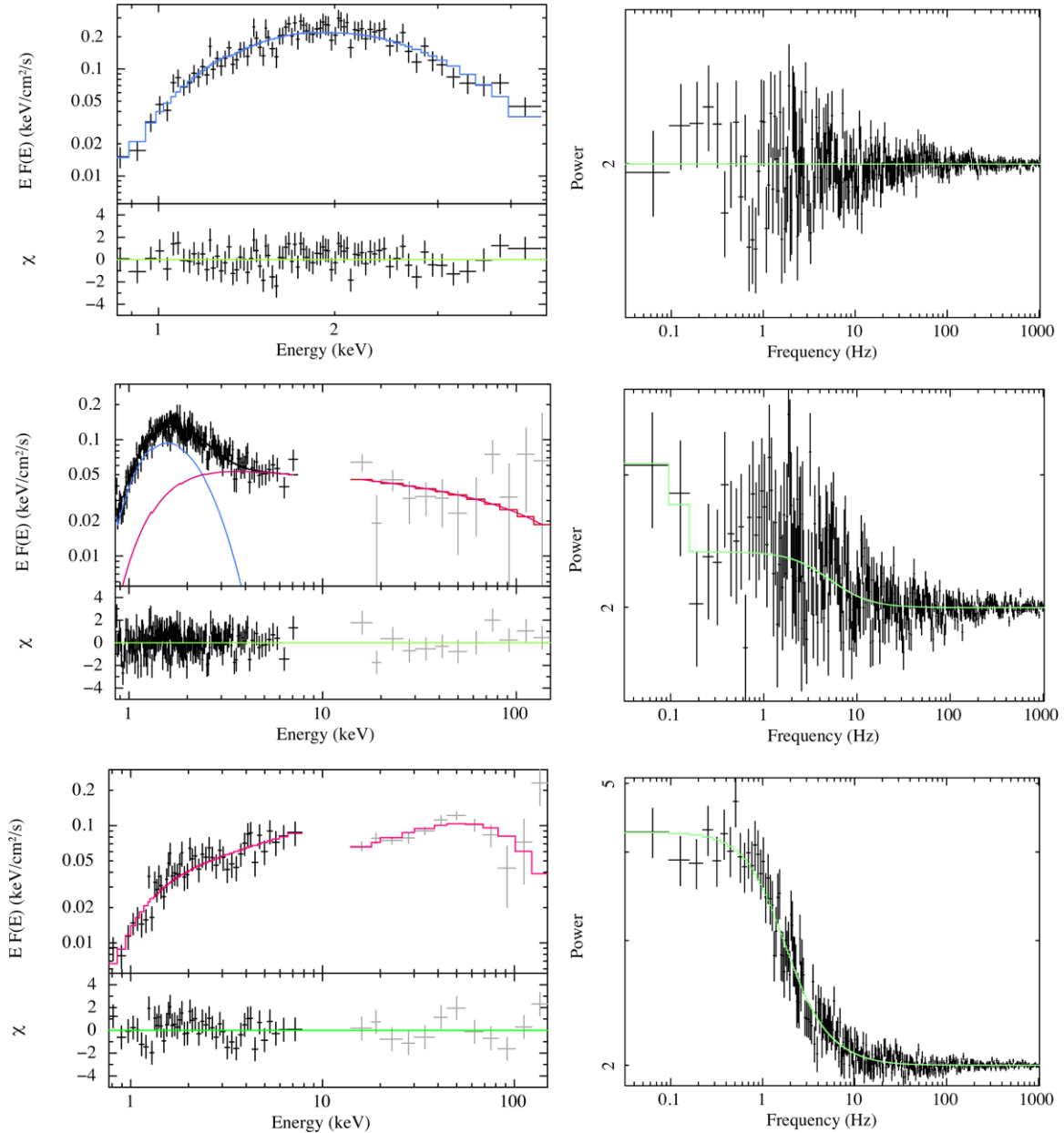


Figure 4. Left-hand panels: *Swift-XRT* (black) and *BAT* (grey) spectra for a sample of three representative observations taken during each spectral state. Right-hand panels: PDS of *NICER* data taken at similar times. The top panels show the X-ray observations taken around MJD 58523 (the source was not detected by *Swift-BAT*), during the soft state, middle panels show the intermediate state (around MJD 58669.76), and the lower panels are during the hard state (observations taken on \sim MJD 58792). Different colours show the different spectral components, where we use blue for *discbb* and magenta for *nthcomp*. Residuals are shown at the bottom of each panel. These three representative observations show a clear change in the X-ray state.

radio position (Fig. 6, lower panel).⁹ Therefore, our radio position suggests that G2 is not associated with MAXI J1810–222. G1, which does not have a *Gaia* parallax distance, is consistent with the *XRT* position but is located ≈ 3.2 -arcsec away from our radio position. From the *Gaia* (Early Data Release 3) catalogue, the probability of finding an unrelated source within 3.2 arcsec from any random position is ≈ 90 per cent. Therefore, we conclude that neither G1

⁹*Gaia* source IDs:

G1 = 4090746312693600512,
G2 = 4090746312736302848.

nor G2 is associated with MAXI J1810–222. In addition, the line-of-sight direction of source is towards the Galactic bulge, so it is not unusual for an optical counterpart to not be detected if the source is sufficiently distant, due to extinction close to the Galactic plane.

4 DISCUSSION

We have studied the broad-band properties of MAXI J1810–222. In this work, we present X-ray spectra of MAXI J1810–222 with *Swift* (*XRT* and *BAT*) and *NuSTAR*, as well as the X-ray timing properties during three representative spectral states with *NICER*. We

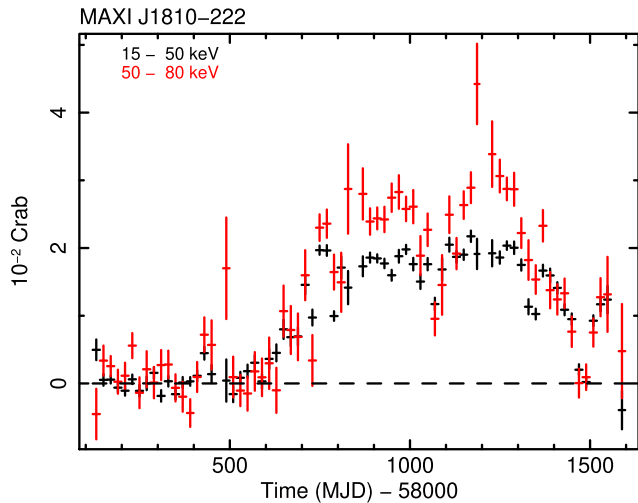


Figure 5. *BAT* light curve showing both the 15- to 50-keV (black) and 50- to 80-keV (red) energy bands. The time binning is 20 d and the X-ray fluxes are reported in units of Crab. Bright flares were observed in the harder energy band between \sim MJD 58850–59000 and \sim MJD 59150–59300, which were not observed in the 15- to 50-keV band.

also provide the result of our coordinated radio campaign with ATCA. While the source was initially detected in the soft state and exhibited a long and complex outburst, the source properties are consistent with an XRB passing through different X-ray spectral states (see e.g. Fender 2006; McClintock & Remillard 2006). During its hard X-ray states the X-ray spectrum displayed an X-ray photon index of ~ 1.6 and a persistent flat-spectrum radio jet. On the contrary, during the soft X-ray states the inner disc emission becomes dominant and the radio emission was either quenched or steep, as expected for a soft state (e.g. Fender et al. 2004a). In addition, we have also observed a number of intermediate states with spectral parameters in between the two main states. Here, we discuss a number of key properties of the system that we can determine from our multiwavelength monitoring campaign.

4.1 Nature of the compact object

Without an estimate on the distance to the source, identifying whether the compact object in an XRB is a BH or NS from radio and X-ray monitoring can be challenging (unless X-ray bursts or pulsations are detected). While our results from MAXI J1810–222 may be able to be reproduced by both a BH and NS accretor, below we argue that the X-ray behaviour of MAXI J1810–222 is more consistent with a BH system.

No X-ray pulsations were detected and no trace of Type-I X-ray bursts were found in the X-ray light curves. If detected, these properties would indicate an NS accretor. However, lack of detection of these two phenomena does not necessarily mean the accretor is not an NS. A large number of the known NS XRBs do not display X-ray pulsations, possibly due to the magnetosphere being absent or very weak (see e.g. Patruno, Wette & Messenger 2018, and references therein) or the polar caps are not favourably aligned. Furthermore, X-ray pulsation behaviour is often intermittent (Casella et al. 2008; Patruno, Altamirano & Messenger 2010; Patruno 2012; Campana, Coti Zelati & D’Avanzo 2013). The exhibited X-ray timing and spectral properties, including the disc temperature and the soft-/hard-state evolution of MAXI J1810–222 are more suggestive of a BH primary. For NS XRBs in their soft states an additional blackbody

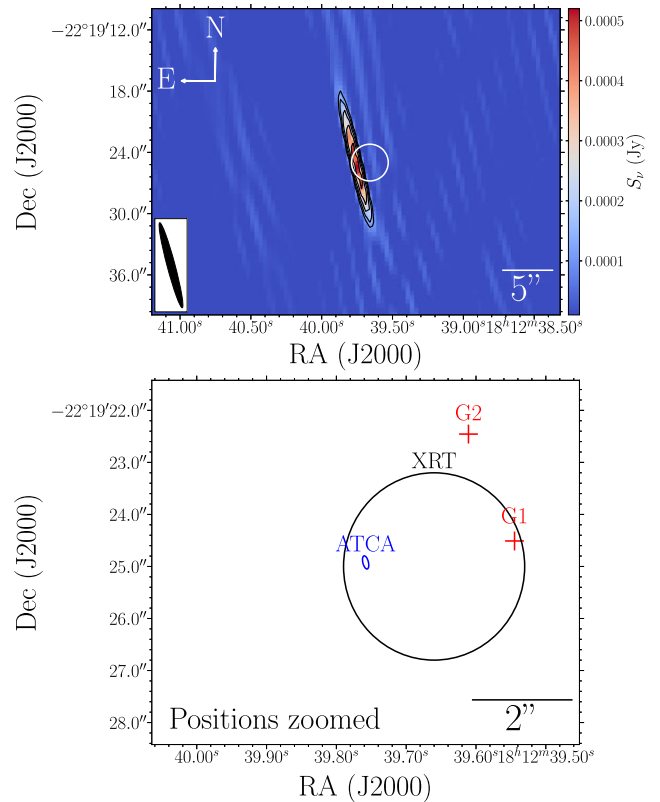


Figure 6. Radio detection and source position of MAXI J1810–222. Top panel: The 9-GHz ATCA image, showing the radio counterpart and the *Swift-XRT* position (error shown by the white circle). The black contour lines are $\sqrt{2}^n$ times the image RMS, where $n = 5, 6, 7, 8, \dots$ and the RMS was $12 \mu\text{Jy beam}^{-1}$. The ATCA beam is shown in the bottom left corner. Bottom panel: A close-in view of the best-fitting ATCA (blue ellipse) and *Swift-XRT* (black circle) source positions, where the errors are shown as the extent of the ellipse/circle. The red crosses (labelled G1 and G2) give the positions of the two nearest *Gaia* sources, where G2 was identified as a possible counterpart to MAXI J1810–222 from *Swift-UVOT* observations (Kennea & Negoro 2019). The errors on the positions of G1 and G2 are on the order of milliarcseconds and are too small to be seen on this figure. Our radio position suggests that neither G1 nor G2 (or any other *Gaia* source) is associated with MAXI J1810–222.

component is generally required to account for the presence of a hot stellar surface. This additional component is typically at a higher temperature (~ 1 – 2 keV) with respect to the disc (see e.g. Barret 2001; D’Ai et al. 2010; Ludlam et al. 2018; Marino et al. 2019). From our monitoring, the highest temperature we derive is ~ 0.6 keV and an additional blackbody component was not required (see e.g. Del Santo et al. 2008).

In addition, even though the estimated values of the X-ray RMS are consistent both with NS and BH systems (Muñoz-Darias et al. 2014), the timing parameters and PDS are more representative of an accreting BH. We do note that the lack of significant QPOs in the PDS may be surprising in the case of a BH binary. However, this could be explained in terms of a low orbital inclination of the source, which would make QPOs seen in the hard and hard intermediate states (i.e. type-C QPOs) much less prominent (Motta et al. 2015). The shape of the PDS continuum and lack of X-ray variability during the soft states are more consistent with a BH accretor (see e.g. Ingram & Motta 2019). Therefore, due to the spectral and timing behaviour, we favour a BH accretor in MAXI J1810–222.

4.2 Outburst evolution and state transitions

In a typical outburst, BH XRBs are usually observed to brighten in a hard X-ray state, before transitioning through the HIMS and SIMS to the soft X-ray state. This hard to soft state transition is usually observed to occur at X-ray luminosities of ≥ 3 per cent of the Eddington luminosity (L_{Edd} ; Dunn et al. 2010). As the outburst decays, the source then typically transitions back to the hard state at X-ray luminosities of between 0.3 per cent and 3 per cent L_{Edd} (Maccarone 2003; Dunn et al. 2010; Vahdat Motlagh, Kalemci & Maccarone 2019). However, some BH XRBs, or specific outbursts from some systems, do not follow this standard pattern of outburst: Sources may not transition beyond the hard or intermediates states (showing no soft states, e.g. Harmon et al. 1994; Brocksopp, Bandyopadhyay & Fender 2004), or can show complex outburst behaviour, such as multiple outburst peaks, rebrightenings, or glitches (e.g. Chen, Shrader & Livio 1997; Homan et al. 2013; Yan & Yu 2017; Parikh et al. 2019; Zhang et al. 2019), and may complete the reverse transition at exceptionally low X-ray luminosities (Tomsick et al. 2014; Chauhan et al. 2019; Russell et al. 2019).

MAXI J1810–222 was initially detected in the soft state (e.g. Maruyama et al. 2018), before transitioning to an apparent long-lived hard state where it remained for ~ 1.5 yr (although we do not have good coverage of the low-energy, soft X-rays over this time, which may variability indicating otherwise). After this long-lived phase, MAXI J1810–222 returned to the soft state, before fading at X-ray wavelengths and transitioning once again to the hard state (Fig. 2). Afterwards, it appears that the source rebrightened again in both *XRT* and *BAT*. However, MAXI J1810–222 was not visible to *XRT* due to the position of the Sun and our broader monitoring campaign had ceased so we did not monitor the source further.

Despite the seemingly complex and unusual outburst pattern, the explanation could be simple if MAXI J1810–222 is located at a sufficiently large distance. In this scenario, the low/hard states typically observed at the beginning of an outburst were simply not bright enough to be detected by all-sky X-ray monitors. The soft→hard transition and subsequent hard state would then simply be a loop/excursion back to a bright hard X-ray state close to the peak of the outburst. Such loops are not unusual in BH XRBs (e.g. Fender et al. 2004a), and sometimes can appear more X-ray bright than the rising hard state (see Tetarenko et al. 2016, and references therein for further discussion). Although, we note that these loops are usually relatively short-lived, lasting weeks and not years like in MAXI J1810–222. This source has remained in outburst for a number of years, and was in a bright harder state for nearly 1.5 yr (however, as mentioned previously, without soft X-ray monitoring we cannot exclude soft X-ray flaring driving state transitions). Long-duration bright hard states are not unheard of; for example, sources such as 4U 0540–697, GRS 1758–258, GRS 1915+105, 4U 1956+350, *Swift* J1753.5–0127, as well as others, have been observed in outburst or within bright hard states for even longer periods of time (see discussions and table 15 in Tetarenko et al. 2016, for a complete list, and references therein). Additionally, the recent outburst of MAXI J1820+070 remained in a bright hard state for more than 100 d during the rise phase of its outburst (e.g. You et al. 2021). Exploring the system properties of BH XRBs that exhibit long-lived hard states does not reveal any obvious shared properties (e.g. Corral-Santana et al. 2016; Tetarenko et al. 2016), with no clear connection between hard-state duration and BH mass, orbital period, inclination, or mass function.

Near the end of our monitoring, MAXI J1810–222 appeared to follow a more standard pattern of outburst decay, where it transitioned

to the soft state, began to fade, and then returned to the hard state. Our final *XRT* observation captured the system brightening once again, which was accompanied by a brightening in *BAT*. This rebrightening was short-lived however, where the *BAT* emission faded once again. Although it is possible the source remained bright in the soft X-ray band. Unfortunately, we were not able to follow the subsequent evolution further with *XRT* due to visibility and telescope constraints, and we did not observe the source again with ATCA.

4.3 Radio emission from MAXI J1810–222

Our radio campaign on MAXI J1810–222 showed radio emission typical of a BH XRB. During the hard states the radio emission was consistent with a flat spectrum, where $\alpha \sim 0$, arising from optically thick synchrotron emission from a self-absorbed compact jet (e.g. Blandford & Königl 1979). During the soft states the radio emission was observed to be either quenched or variable with a steep radio spectrum, indicative of an optically-thin transient jet that is launched around the hard to soft state transition (e.g. Fender 2006).

We also note that the flat spectrum remained throughout the soft X-ray brightening (\sim MJD 59322), suggesting that while the X-ray spectrum did soften, the steady jet remained on, suggesting that the source likely only transitioned as far as an intermediate state.

4.3.1 Location on the radio/X-ray plane

Placing our quasi-simultaneous hard-state radio and X-ray observations of MAXI J1810–222 on the radio/X-ray plane allows us to compare its inferred radio and X-ray luminosities against typical luminosities of BH (and NS) XRBs (Fig. 7). Due to the unknown source distance, the observed radio and X-ray luminosities are consistent with those typically observed from both NS and BH XRBs (and we caution against using this method as the sole method to identify the nature of the compact object). However, as discussed previously, the observed X-ray spectral and timing properties suggest that MAXI J1810–222 is most likely a BH XRB. If that is the case, comparing its hard-state radio and X-ray luminosities against those typically observed from other BH systems, to be comparable in luminosity MAXI J1810–222 needs to be located at a distance of $\gtrsim 6$ kpc. We note that some systems deviate from the standard radio/X-ray luminosities shown by the broad population of BH XRBs, displaying a hybrid correlation and occupying a different parameter space (e.g. MAXI J1348–630; Carotenuto et al. 2021b). Therefore, we caution against its use as anything more than a suggestion that MAXI J1810–222 is relatively distant, although that proposition is also supported by the X-ray non-detection during the rising hard state (by *Swift-BAT* and MAXI).

Out of interest, we also show the line-of-sight direction to the source against a top-down view of the Milky Way showing that MAXI J1810–222 can still be located in the Galactic Bar region out to ~ 12 kpc (Fig. 8), implying larger distances remain plausible while still residing within higher density Galactic regions (which does not begin until ~ 6 kpc in that direction). Using the Galactic mass density model of low-mass XRBs (Grimm, Gilfanov & Sunyaev 2002) as implemented in Atri et al. (2019), we estimate the probability of the source being at distances above 6 kpc to be 94 per cent (and 31 per cent at > 10 kpc).

In support of a distance to the source in excess of 6 kpc, our *Swift-XRT* gives an N_{H} , which traces the interstellar gas along the line of sight, of $\sim 1 \times 10^{22} \text{ cm}^{-2}$. This N_{H} is more than twice the expected Galactic contribution (with a weighted average of $4.2 \times 10^{21} \text{ cm}^{-2}$

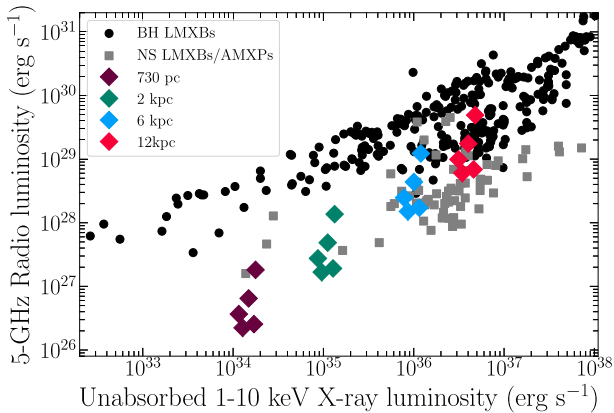


Figure 7. The radio and X-ray luminosities of MAXI J1810–222 for varying source distances (diamonds), plotted with the full sample of accreting BHs (black circles) and NS (grey squares), where often multiple observations are plotted for a single source at different luminosities (data taken from Bahramian et al. 2018). For completeness, we also include the sample of accreting millisecond X-ray pulsars and transitional millisecond pulsars within the NS sample. We only show quasi-simultaneous data taken when the source is securely in the hard X-ray state. If MAXI J1810–222 is indeed a BH XRB, for the observed luminosities to be consistent with other BH systems, the source would need to be located at a distance in excess of 6 kpc.

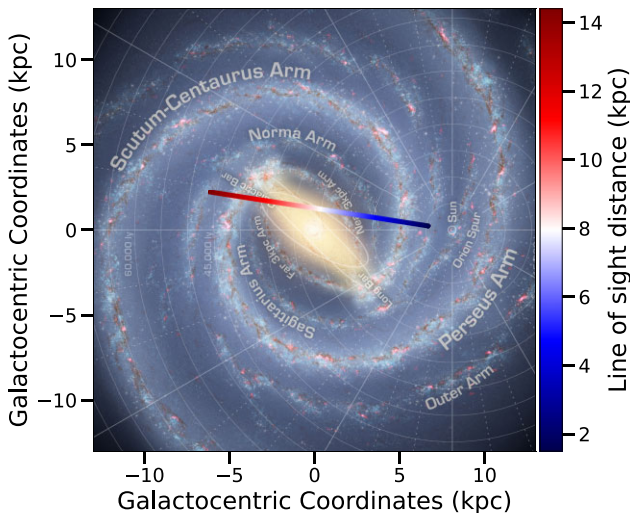


Figure 8. The line of sight towards of MAXI J1810–222 plotted against the Milky Way in Galactocentric coordinates. The coloured line represents the direction to MAXI J1810–222 for different distances to the source (as shown by the colour bar), indicating that the system can lie at large distances while remaining in the denser regions of the Galactic centre. The background image is an illustration of the Milky Way (illustration credit: NASA/JPL-Caltech/ESO/R. Hurt), image created using the python package MW-PLOT.

being expected) along the line of sight (HI4PI Collaboration et al. 2016). While a higher than expected absorption along the line of sight does not indicate the true distance to a source, it is generally thought to be a result of a source’s close proximity to the Galactic centre (e.g. Miller et al. 2006). High resolution three-dimensional optical reddening maps suggest that the majority of the Galactic extinction along the line of sight to MAXI J1810–222 lies between 4.5 and 6 kpc, suggestive that MAXI J1810–222 lies beyond this distance. Although, higher than expected extinction can also be a result of the source being at such a high inclination that the disc

material or donor star passes through the line of sight (e.g. Parmar et al. 2002). However, no X-ray dips are detected in the extensive X-ray monitoring of this source implying this is likely not the case. Additionally, the lack of any QPO detection is more reminiscent of a low-inclination source. As such, the measured N_{H} and line-of-sight direction is suggestive that MAXI J1810–222 lies at larger distances, closer to the Galactic centre where the higher than expected absorption may arise from within the Galactic bulge.

During its hard X-ray states, MAXI J1810–222 showed a large variation in radio luminosity in comparison to its X-ray evolution, where the radio luminosity was observed to change by a factor of ~ 10 and the X-ray only by a factor of ~ 2 . Such large radio variation with respect to the X-rays is not regularly observed in BH systems. However, it has been observed in at least one low-inclination BH XRB, where large variations in the radio emission could be explained by variable Doppler boosting of the jet emission and a low inclination (Russell et al. 2015). As shown in Fig. 5, it is possible that the larger than expected variations in the hard-state radio flux densities may be related to flaring observed in the hardest 50- to 80-keV X-rays, which are not captured in the standard 1- to 10-keV band used for typical radio/X-ray correlations. Speculatively, there may be a suggestion of an anti-correlation between the 50- to 80-keV flux and the radio flux density during the brightest hard state, where our brightest radio detections occurred during the fainter 50- to 80-keV periods. Connecting the radio flux to the hardest X-ray emission could help to understand the origin of the additional and variable soft gamma-ray component observed by the *INTEGRAL* X-ray telescope in several BH XRBs during their hard states (see e.g. Del Santo et al. 2008, Droulans et al. 2010). However, between \sim MJD 59150 and 59300 the 50- to 80-keV X-rays were bright and exhibiting strong variability. Over the same period, no significant radio flaring was observed, with the radio emission remaining faint and steady. If there was a simple and direct connection between the hard X-rays and the jet, we would naively expect similar changes in the radio emission. Hence, it is likely that such a connection, if there, is more complex and not observed within our data. As such, any connection between the jet and additional (to the thermal Comptonization) hard X-ray component is speculative. A combination of multiwavelength observations and physical models would be necessary for such a study (e.g. Bassi et al. 2020).

5 CONCLUSIONS

MAXI J1810–222 is a Galactic X-ray transient that was first detected in outburst in 2019. X-ray spectral and timing properties are most consistent with it being BH XRB, although an NS counterpart cannot be ruled out. At first glance, this source has displayed an odd outburst evolution, first being detected in a soft state, before transitioning to a bright hard state where it remained for ~ 1.5 yr. After which, the source returned to the soft state, faded and then transitioned to a hard state as the outburst apparently decayed. However, the seemingly odd outburst evolution can be simply explained with a sufficiently large distance to the source, such that the source was not detectable by all sky X-ray monitors (which primarily work in the harder X-ray bands) during its lower luminosity rise phase. In this scenario, the initial bright soft \rightarrow hard state transition could be a result of a loop/excursion back to the hard state from the high soft state, and not a standard reverse transition during the outburst decay. Such excursions are not rare in BH systems. At later times, the source appeared to follow a more standard X-ray evolution. Our radio detection also rules out the proposed association with a proposed *Gaia* counterpart (Kennea & Negoro 2019).

ACKNOWLEDGEMENTS

We would like to thank the anonymous referee for their helpful comments. Authors at INAF acknowledge financial contribution from the agreement ASI-INAF n.2017-14-H.0 and INAF mainstream. This work has been partially supported by the ASI-INAF program I/004/I1/5. The Australia Telescope Compact Array is part of the Australia Telescope National Facility which is funded by the Australian Government for operation as a National Facility managed by CSIRO. We acknowledge the Gomeri people as the traditional owners of the observatory site. This research has made use of (i) NASA's Astrophysics Data System, (ii) data, software, and/or web tools obtained from the High Energy Astrophysics Science Archive Research Center (HEASARC), a service of the Astrophysics Science Division at NASA Goddard Space Flight Center (GSFC) and of the Smithsonian Astrophysical Observatory's High Energy Astrophysics Division, (iii) data supplied by the UK *Swift* Science Data Centre at the University of Leicester, and (iv) MAXI data provided by RIKEN, JAXA, and the MAXI team. This work has made use of data from the European Space Agency (ESA) mission *Gaia* (<https://www.cosmos.esa.int/gaia>), processed by the *Gaia* Data Processing and Analysis Consortium (DPAC, <https://www.cosmos.esa.int/web/gaia/dpac/consortium>). Funding for the DPAC has been provided by national institutions, in particular the institutions participating in the *Gaia* Multilateral Agreement. AM is supported by the H2020 European Research Council Consolidator Grant 'MAGNESIA' under grant agreement No. 817661 (PI: Rea) and National Spanish Grant PGC2018-095512-BI00. This work was also partially supported by the program Unidad de Excelencia Maria de Maeztu CEX2020-001058-M, and by the PHAROS COST Action (No. CA16214).

DATA AVAILABILITY

Data from *Swift* are publicly available from HEASARC (<https://heasarc.gsfc.nasa.gov/>). Best-fitting X-ray parameters are provided in the manuscript. Raw ATCA data are provided on the Australia Telescope Online Archive (<https://atoa.atnf.csiro.au/query.jsp>), under project code CX445. All calibrated radio flux densities used in this work are provided.

REFERENCES

Atri P. et al., 2019, *MNRAS*, 489, 3116
 Bahramian A. et al., 2018, Radio/X-ray correlation database for X-ray binaries, <https://doi.org/10.5281/zenodo.1252036>
 Bailer-Jones C. A. L., Rybizki J., Fousneau M., Mantelet G., Andrae R., 2018, *AJ*, 156, 58
 Barret D., 2001, *Adv. Space Res.*, 28, 307
 Bassi T. et al., 2020, *MNRAS*, 494, 571
 Belloni T. M., Motta S. E., 2016, in Bambi C., ed., *Astrophysics and Space Science Library*, Vol. 440, *Astrophysics of Black Holes: From Fundamental Aspects to Latest Developments*. Springer, Berlin, p. 61
 Blandford R. D., Königl A., 1979, *ApJ*, 232, 34
 Bright J. S. et al., 2020, *Nature Astron.*, 4, 697
 Brocksopp C., Bandyopadhyay R. M., Fender R. P., 2004, *New Astron.*, 9, 249
 Campana S., Coti Zelati F., D'Avanzo P., 2013, *MNRAS*, 432, 1695
 Carotenuto F., Tremou E., Corbel S., Fender R., Woudt P., Miller-Jones J., 2019, *Astron. Telegram*, 12521, 1
 Carotenuto F. et al., 2021a, *MNRAS*, 504, 444
 Carotenuto F. et al., 2021b, *MNRAS*, 505, L58
 Casella P., Altamirano D., Patruno A., Wijnands R., van der Klis M., 2008, *ApJ*, 674, L41

Chauhan J. et al., 2019, *MNRAS*, 488, L129
 Chen W., Shrader C. R., Livio M., 1997, *ApJ*, 491, 312
 Corbel S., Fender R. P., 2002, *ApJ*, 573, L35
 Corbel S., Fender R. P., Tzioumis A. K., Nowak M., McIntyre V., Durouchoux P., Sood R., 2000, *A&A*, 359, 251
 Corbel S., Fender R. P., Tzioumis A. K., Tomsick J. A., Orosz J. A., Miller J. M., Wijnands R., Kaaret P., 2002, *Science*, 298, 196
 Corbel S., Nowak M. A., Fender R. P., Tzioumis A. K., Markoff S., 2003, *A&A*, 400, 1007
 Corbel S., Fender R. P., Tomsick J. A., Tzioumis A. K., Tingay S., 2004, *ApJ*, 617, 1272
 Corbel S., Coriat M., Brocksopp C., Tzioumis A. K., Fender R. P., Tomsick J. A., Buxton M. M., Bailyn C. D., 2013a, *MNRAS*, 428, 2500
 Corbel S. et al., 2013b, *MNRAS*, 431, L107
 Coriat M. et al., 2011, *MNRAS*, 414, 677
 Corral-Santana J. M., Casares J., Muñoz-Darias T., Bauer F. E., Martínez-Pais I. G., Russell D. M., 2016, *A&A*, 587, A61
 D'Ài A. et al., 2010, *A&A*, 516, A36
 Del Santo M., Malzac J., Jourdain E., Belloni T., Ubertini P., 2008, *MNRAS*, 390, 227
 Díaz Trigo M., Boirin L., 2016, *Astron. Nachr.*, 337, 368
 Done C., Gierliński M., Kubota A., 2007, *A&AR*, 15, 1
 Droulans R., Belmont R., Malzac J., Jourdain E., 2010, *ApJ*, 717, 1022
 Ducci L. et al., 2020, *Astron. Telegram*, 13540, 1
 Dunn R. J. H., Fender R. P., Körding E. G., Belloni T., Cabanac C., 2010, *MNRAS*, 403, 61
 Fender R., 2006, *Jets From X-ray Binaries*. Cambridge Univ. Press, Cambridge, p. 381
 Fender R., 2010, *Lecture Notes in Physics*, Vol. 794, *The Jet Paradigm*. Springer, Berlin, p. 115
 Fender R. P., 2001, *MNRAS*, 322, 31
 Fender R. P., Kuulkers E., 2001, *MNRAS*, 324, 923
 Fender R. P., Belloni T. M., Gallo E., 2004a, *MNRAS*, 355, 1105
 Fender R., Wu K., Johnston H., Tzioumis T., Jonker P., Spencer R., van der Klis M., 2004b, *Nature*, 427, 222
 Fender R. P., Homan J., Belloni T. M., 2009, *MNRAS*, 396, 1370
 Fomalont E. B., Geldzahler B. J., Bradshaw C. F., 2001, *ApJ*, 553, L27
 Gaia Collaboration, 2016, *A&A*, 595, A1
 Gaia Collaboration, 2018, *A&A*, 616, A1
 Gallo E., Fender R. P., Pooley G. G., 2003, *MNRAS*, 344, 60
 Gallo E., Miller B. P., Fender R., 2012, *MNRAS*, 423, 590
 Gallo E., Degenaar N., van den Eijnden J., 2018, *MNRAS*, 478, L132
 Galloway D. K., Keek L., 2021, in Belloni T. M., Méndez M., Zhang C., eds, *Astrophysics and Space Science Library*, Vol. 461, *Timing Neutron Stars: Pulsations, Oscillations and Explosions*. Springer, Berlin, p. 209
 Grimm H. J., Gilfanov M., Sunyaev R., 2002, *A&A*, 391, 923
 Gusinskaia N. V. et al., 2017, *MNRAS*, 470, 1871
 Gusinskaia N. V. et al., 2020, *MNRAS*, 492, 2858
 Hannikainen D. C., Hunstead R. W., Campbell-Wilson D., Sood R. K., 1998, *A&A*, 337, 460
 Harmon B. A. et al., 1994, *ApJ*, 425, L17
 HI4PI Collaboration et al., 2016, *A&A*, 594, A116
 Homan J., Fridriksson J. K., Jonker P. G., Russell D. M., Gallo E., Kuulkers E., Rea N., Altamirano D., 2013, *ApJ*, 775, 9
 Ingram A. R., Motta S. E., 2019, *New Astron. Rev.*, 85, 101524
 Jamil O., Fender R. P., Kaiser C. R., 2010, *MNRAS*, 401, 394
 Kalemci E., Dinger T., Tomsick J. A., Buxton M. M., Bailyn C. D., Chun Y. Y., 2013, *ApJ*, 779, 95
 Kennea J. A., Negro H., 2019, *Astron. Telegram*, 12487, 1
 Kuulkers E., den Hartog P. R., in 't Zand J. J. M., Verbunt F. W. M., Harris W. E., Cocchi M., 2003, *A&A*, 399, 663
 Lasota J.-P., 2001, *New Astron. Rev.*, 45, 449
 Lewin W. H. G., van Paradijs J., Taam R. E., 1993, *Space Sci. Rev.*, 62, 223
 Ludlam R. M. et al., 2018, *ApJ*, 858, L5
 Maccarone T. J., 2003, *A&A*, 409, 697
 Maccarone T. J., 2008, in Evans A., Bode M. F., O'Brien T. J., Darnley M. J., eds, *ASP Conf. Ser. Vol. 401, RS Ophiuchi (2006) and the Recurrent Nova Phenomenon*. Astron. Soc. Pac., San Francisco, p. 191

- McClintock J. E., Remillard R. A., 2006, *Black Hole Binaries*. Cambridge Univ. Press, Cambridge, p. 157
- McConnell D., Sadler E. M., Murphy T., Ekers R. D., 2012, *MNRAS*, 422, 1527
- McMullin J. P., Waters B., Schiebel D., Young W., Golap K., 2007, in Shaw R. A., Hill F., Bell D. J., eds, *ASP Conf. Ser. Vol. 376, Astronomical Data Analysis Software and Systems XVI*. Astron. Soc. Pac., San Francisco, p. 127
- Marino A. et al., 2019, *MNRAS*, 490, 2300
- Marino A. et al., 2020, *MNRAS*, 498, 3351
- Maruyama W. et al., 2018, *Astron. Telegram*, 12264, 1
- Massardi M., Bonaldi A., Bonavera L., López-Caniego M., de Zotti G., Ekers R. D., 2011, *MNRAS*, 415, 1597
- Matsuoka M. et al., 2009, *PASJ*, 61, 999
- Migliari S., 2011, in Romero G. E., Sunyaev R. A., Belloni T., eds, *Proc. IAU Symp. 275, Jets at All Scales*. Kluwer, Dordrecht, p. 233
- Migliari S., Fender R. P., 2006, *MNRAS*, 366, 79
- Migliari S., Fender R. P., Rupen M., Jonker P. G., Klein-Wolt M., Hjellming R. M., van der Klis M., 2003, *MNRAS*, 342, L67
- Migliari S., Fender R. P., Rupen M., Wächter S., Jonker P. G., Homan J., van der Klis M., 2004, *MNRAS*, 351, 186
- Miller J. M. et al., 2006, *ApJ*, 646, 394
- Miller-Jones J. C. A. et al., 2010, *ApJ*, 716, L109
- Miller-Jones J. C. A. et al., 2012, *MNRAS*, 421, 468
- Motta S. E., Casella P., Henze M., Muñoz-Darias T., Sanna A., Fender R., Belloni T., 2015, *MNRAS*, 447, 2059
- Muñoz-Darias T., Fender R. P., Motta S. E., Belloni T. M., 2014, *MNRAS*, 443, 3270
- Negoro H. et al., 2018, *Astron. Telegram*, 12254, 1
- Negoro H. et al., 2019, *Astron. Telegram*, 12910, 1
- Oeda M. et al., 2019, *Astron. Telegram*, 12398, 1
- Parikh A. S., Russell T. D., Wijnands R., Miller-Jones J. C. A., Sivakoff G. R., Tetarenko A. J., 2019, *ApJ*, 878, L28
- Parmar A. N., Oosterbroek T., Boirin L., Lumb D., 2002, *A&A*, 386, 910
- Patruno A., 2012, *ApJ*, 753, L12
- Patruno A., Altamirano D., Messenger C., 2010, *MNRAS*, 403, 1426
- Patruno A., Wette K., Messenger C., 2018, *ApJ*, 859, 112
- Pringle J. E., Rees M. J., Pacholczyk A. G., 1973, *A&A*, 29, 179
- Rushton A. P. et al., 2017, *MNRAS*, 468, 2788
- Russell T. D., Soria R., Miller-Jones J. C. A., Curran P. A., Markoff S., Russell D. M., Sivakoff G. R., 2014, *MNRAS*, 439, 1390
- Russell T. D. et al., 2015, *MNRAS*, 450, 1745
- Russell T. D., Degenaar N., Wijnands R., van den Eijnden J., Gusinskaia N. V., Hessels J. W. T., Miller-Jones J. C. A., 2018, *ApJ*, 869, L16
- Russell T. D. et al., 2019, *ApJ*, 883, 198
- Russell T. D. et al., 2020, *MNRAS*, 498, 5772
- Rutledge R., Moore C., Fox D., Lewin W., van Paradijs J., 1998, *Astron. Telegram*, 8, 1
- Segreto A., Cusumano G., Ferrigno C., La Parola V., Mangano V., Mineo T., Romano P., 2010, *A&A*, 510, A47
- Shakura N. I., Sunyaev R. A., 1973, *A&A*, 24, 337
- Tetarenko B. E., Sivakoff G. R., Heinke C. O., Gladstone J. C., 2016, *ApJS*, 222, 15
- Tetarenko A. J. et al., 2017, *MNRAS*, 469, 3141
- Tomsick J. A., Yamaoka K., Corbel S., Kalemci E., Migliari S., Kaaret P., 2014, *ApJ*, 791, 70
- Tudor V. et al., 2017, *MNRAS*, 470, 324
- Tudose V., Fender R. P., Linares M., Maitra D., van der Klis M., 2009, *MNRAS*, 400, 2111
- Vahdat Motlagh A., Kalemci E., Maccarone T. J., 2019, *MNRAS*, 485, 2744
- van den Eijnden J. et al., 2021, *MNRAS*, 507, 3899
- van Paradijs J., Verbunt F., 1984, in Woosley S. E., ed., *AIP Conf. Proc. Vol. 115, High Energy Transients in Astrophysics*. Am. Inst. Phys., New York, p. 49
- Verner D. A., Ferland G. J., Korista K. T., Yakovlev D. G., 1996, *ApJ*, 465, 487
- Wilms J., Allen A., McCray R., 2000, *ApJ*, 542, 914
- Wood C. M. et al., 2021, *MNRAS*, 505, 3393
- Yan Z., Yu W., 2017, *MNRAS*, 470, 4298
- You B. et al., 2021, *Nature Commun.*, 12, 1025
- Zdziarski A. A., Gierliński M., 2004, *Prog. Theor. Phys. Suppl.*, 155, 99
- Zhang G. B. et al., 2019, *ApJ*, 876, 5

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.