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Comparison of hydrostatic and lensing cluster mass estimates: A pilot study in MACS J0647.7+7015

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ABSTRACT

The detailed characterization of scaling laws relating the observables of a cluster of galaxies to their mass is crucial for obtaining accurate cosmological constraints with clusters. In this paper, we present a comparison between the hydrostatic and lensing mass profiles of the cluster MACS J0647.7+7015 at $z = 0.59$. The hydrostatic mass profile is obtained from the combination of high resolution NIKA2 thermal Sunyaev-Zel'dovich and *XMM-Newton* X-ray observations of the cluster. The lensing mass profile, on the other hand, is obtained from an analysis of the CLASH lensing data based on the lensing convergence map. We find significant variation in the cluster mass estimate depending on the observable, the modeling of the data, and the knowledge of the cluster's dynamical state. This might lead to significant systematic effects on cluster cosmological analyses for which only a single observable is generally used. From this pilot study, we conclude that the combination of high resolution Sunyaev-Zel'dovich, X-ray, and lensing data could allow us to identify and correct for these systematic effects. This would constitute a very interesting extension of the NIKA2 SZ Large Program.

Key words. gravitational lensing: weak – gravitational lensing: strong – cosmology: observations – X-rays: galaxies: clusters – large-scale structure of Universe

1. Introduction

Clusters of galaxies are the largest gravitationally bound objects in the Universe and constitute the last step of the hierarchical process of structure formation (see [Kravtsov & Borgani 2012](#), for a review). Therefore, their abundance in mass and redshift and their spatial distribution are powerful cosmological probes. Cluster cosmology is particularly sensitive to the primordial density fluctuations and the expansion history and matter content of the Universe (see [Allen et al. 2011](#), for a review). Clusters are primarily composed of a dark matter halo, hot baryonic gas, and individual galaxies, which correspond to about 85, 12, and 3% of their total mass, respectively. Thanks to this multi-component nature, the detection and study of clusters of galaxies can be performed via a large number of complementary observables across wavelengths: optical and IR emission from the galaxies in the cluster ([Bahcall 1977](#)), gravitational lensing effects from radio to optical wavelengths (see [Bartelmann 2010](#), for a review), X-ray emission from the hot baryonic gas ([Sarazin 1988](#); [Böhringer & Werner 2010](#)), and the thermal Sunyaev-Zel'dovich (tSZ) effect at microwave and millimeter wavelengths ([Sunyaev & Zel'dovich 1972, 1980](#)).

In the last decade, large catalogues of clusters of galaxies have been made available at different wavelengths, leading to a large

number of cosmological studies (e.g., [Planck Collaboration XX 2014](#); [Planck Collaboration XXI 2014](#); [Böhringer & Chon 2015](#); [de Haan et al. 2016](#); [Böhringer et al. 2017](#); [Pacaud et al. 2018](#); [Bocquet et al. 2019](#); [Costanzi et al. 2019](#); [Abbott et al. 2020](#)). From these studies, which are based mainly on cluster number counts as a function of mass and redshift, it can be concluded that cosmological parameter constraints from galaxy clusters may be affected by systematic uncertainties. These uncertainties come from both the theoretical and observational modeling of clusters (see [Salvati et al. 2020](#), for a recent summary).

On the one hand, the modeling of the halo mass function from numerical simulations is not unique and may lead to uncertainties of about 10% on the final cosmological parameters. On the other hand, cluster masses are inferred from cluster observables, either directly or via scaling relations. These mass estimates may both be affected by observational and modeling statistical and systematic uncertainties (see [Pratt et al. 2019](#); [Ruppin et al. 2019a](#), for a review). The *Planck* 2013 and 2015 results have allowed for a direct comparison between cluster-based cosmology and cosmology based on the cosmic microwave background (CMB); discrepancy between the two has been shown at about the 2σ level ([Planck Collaboration XX 2014](#); [Planck Collaboration XXI 2014](#); [Planck Collaboration XXIV 2016](#); [Planck Collaboration XXII 2016](#); [Bolliet et al. 2018](#)),

which was slightly reduced for the *Planck* 2018 CMB results (Planck Collaboration I 2020). From the cluster analysis point of view, likely explanations for this discrepancy are deviations from the self-similarity and hydrostatic equilibrium (HE) assumptions at the origin of the scaling relations that link the total mass of the cluster to the cluster tSZ emission (Planck Collaboration Int. XI 2013; Planck Collaboration Int. III 2013). For example, one would expect redshift evolution of the scaling laws due to the cluster's dynamical state and variations in the environmental conditions (e.g., Salvati et al. 2019). Such evolution has generally not been taken into account to date as most scaling laws have been derived from low redshift clusters (e.g., Planck Collaboration Int. III 2013). In particular, cluster dynamical state variations induce significant variations (see Gianfagna et al. 2021, and references therein) in the hydrostatic mass bias, which relates the hydrostatic mass to the total mass of the cluster. For the most disturbed clusters, the self-similarity approximation may not be valid.

The identification of the dynamical state of clusters requires resolved observations as well as realistic simulations of clusters so that robust indicators can be identified (De Petris et al. 2020). Thermal Sunyaev-Zel'dovich and X-ray cluster observations can be used to identify inhomogeneities in the intracluster medium (ICM), such as over-pressure or high temperature regions as well as shocks. Equivalently, the spatial distribution and velocity of the galaxies in the cluster as well as the gravitational lensing estimates of the cluster gravitational potential can be used to identify substructures. Moreover, these different observables will lead to independent cluster mass estimates with different systematic uncertainties (see Pratt et al. 2019, for a detailed description).

As a consequence, to extend scaling laws to high redshift clusters, there is a need for high resolution observations of a representative sample of clusters at different wavelengths. This will be the case for the NIKA2 SZ Large Program (LPSZ) sample that is composed of 45 *Planck* and ACT SZ-selected clusters in the $0.5 < z < 0.9$ redshift range (Mayet et al. 2020; Macias-Pérez et al. 2017). NIKA2 is a new-generation continuum camera installed at the IRAM 30 m single-dish telescope (Adam et al. 2018; Calvo et al. 2016). The combination of a large field of view (6.5 arcmin), a high angular resolution (17.7 arcsec at 150 GHz), and a high sensitivity of $8 \text{ mJy s}^{1/2}$ at 150 GHz provides the NIKA2 camera with unique tSZ mapping capabilities (Perotto et al. 2020). High resolution X-ray maps of the NIKA2 Large Program clusters are also being obtained using the *XMM-Newton* satellite. Furthermore, the LPSZ will be extended with complementary optical and IR data, which are expected to be obtained from archive data and/or on-purpose observations. Moreover, a NIKA2 LPSZ twin sample composed of synthetic clusters extracted from a hydrodynamical simulation data set are also available (Ruppin et al. 2019b).

As a demonstration of the interest of these multiwavelength observations, we present in this paper a comparison between the hydrostatic and lensing mass profiles of the cluster MACS J0647.7+7015, also known as PSZ2 G144.83+25.11. The former is obtained from the combination of low resolution *Planck* tSZ, high resolution NIKA2 SZ, and *XMM-Newton* X-ray observations of the cluster (Ruppin et al. 2018). The latter is obtained from a reanalysis of the CLASH lensing data (Zitrin et al. 2015). The paper is organized as follows. Section 2 presents the different hydrostatic mass estimates used in this work. In Sect. 3 we discuss the estimation of the lensing mass profile. The comparison of the different mass profile estimates is presented in Sect. 4. Finally, we draw conclusions in Sect. 5.

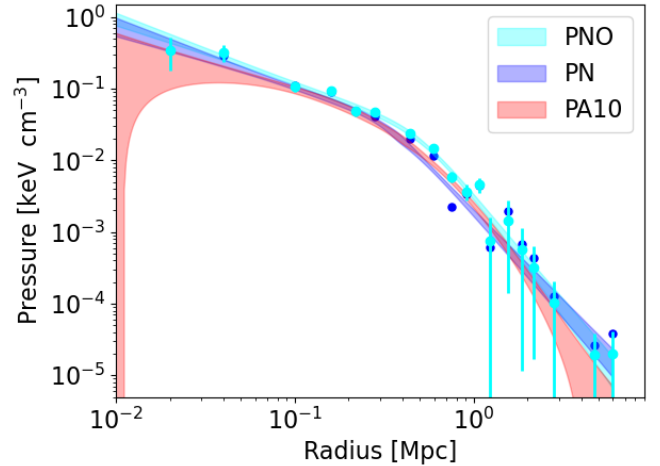


Fig. 1. Pressure profiles derived from the full 150 GHz map (cyan, PNO, cyan) and excluding the over-pressure region (PN, blue). The dots and error bars correspond to the non-parametric fit to the SZ. The shadow region represents the best-fit gNFW model to the non-parametric pressure profiles. We also show in red the universal pressure profile (Arnaud et al. 2010) for the cluster (PA10).

2. Hydrostatic mass profile

In this section we discuss and compare different hydrostatic mass radial profiles for the MACS J0647.7+7015 cluster. They were determined from various estimates of the ICM pressure, temperature, and electron density radial profiles obtained using the *Planck* and NIKA2 tSZ data and the *XMM-Newton* X-ray observations as presented in Ruppin et al. (2018). Here, we briefly describe the most important steps of the analysis and present extra material not included in Ruppin et al. (2018).

2.1. tSZ-based pressure profile estimates

The tSZ-based pressure profile estimates for the ICM of MACS J0647.7+7015 were obtained by Ruppin et al. (2018) from a combination of the NIKA2 cluster surface brightness map at 150 GHz (Ruppin et al. 2018) and the *Planck* Compton parameter map (Planck Collaboration XXII 2016). From the NIKA2 brightness map it was found that the tSZ signal exhibits an elliptical morphology with a major axis oriented along the east-west direction and a thermal pressure excess in the southwest region. As this over-pressure region might have a significant impact on the determination of the hydrostatic mass of the cluster, two different thermal pressure profile estimates were considered, by excluding or not excluding the over-pressure region. In both cases, the thermal pressure profiles were extracted by Ruppin et al. (2018) via a Monte Carlo Markov chain (MCMC) analysis, accounting for instrumental properties. Following the work of Ruppin et al. (2017), the deprojected cluster pressure profile was constrained non-parametrically from the cluster core ($R \sim 0.02 R_{500}$) to its outskirts ($R \sim 3 R_{500}$). The best-fit non-parametric pressure profile was also fitted to a parametric generalized Navarro-Frenk-White (gNFW) model (Nagai et al. 2007) in order to be able to extrapolate at large cluster-centric radii. The derived pressure profiles and best-fit gNFW models are presented in Fig. 1. We observe that the pressure profiles are well measured up to 9 Mpc from the center of the cluster, which is well beyond the expected R_{500} radius for the cluster.

In summary, we consider in this work three tSZ-based estimates of the ICM thermal pressure profile of

MACSJ0647.7+7015: (1) a NIKA2-*Planck* combined analysis including the over-pressure region (PNO), (2) a NIKA2-*Planck* combined analysis excluding the over-pressure region (PN), and (3): the universal pressure profile (PA10; Arnaud et al. 2010). The last, which is also presented in Fig. 1, was obtained by fitting a gNFW model to the data with the slope parameters a , b , and c fixed to the universal pressure profile values (Arnaud et al. 2010) and P_0 and r_p as free parameters.

2.2. X-ray-based electron density and pressure profiles

Deep X-ray observations of MACSJ0647.7+7015 by *XMM-Newton* (effective exposure time of ~ 68 ks) permit the reconstruction of both the electron density and temperature profiles, from which the pressure profile can also be determined. We used the data presented in Ruppin et al. (2018) that were processed using the standard procedures presented, for example, in Bartalucci et al. (2017). We notice that, as shown by Ruppin et al. (2018), no significant over-density region was identified in the southwest region of MACSJ0647.7+7015 in the *XMM-Newton* X-ray surface brightness map. However, we considered two estimates of the electron density radial profile obtained by excluding or not the over-pressure region. The former was already presented in Ruppin et al. (2018). In both cases, the same procedure was used to constrain the gas density profile from the *XMM-Newton* observations, and it is described in detail in Pratt et al. (2010) and Planck Collaboration Int. III (2013).

The deprojected gas density profiles were fitted by a simplified Vikhlinin et al. (2006) parametric model given by

$$n_e(r) = n_{e0} \left[1 + \left(\frac{r}{r_c} \right)^2 \right]^{-3\beta/2} \left[1 + \left(\frac{r}{r_s} \right)^\gamma \right]^{-\epsilon/2\gamma}, \quad (1)$$

where n_{e0} is the central gas density, r_c is the core radius, and r_s is the transition radius at which an additional steepening in the profile occurs. The β and ϵ parameters define the inner and the outer profile slopes, respectively. The γ parameter gives the width of the transition in the profile. The value of the γ parameter was fixed to 3 since it provides a good description of all the clusters considered in the analysis of Vikhlinin (2006).

The *XMM-Newton* data can also be used to estimate the cluster pressure profile by combining the gas density profiles discussed above and the gas temperature determined from the spectroscopic observations via a deprojection procedure as presented in Pratt et al. (2010) and Planck Collaboration Int. III (2013). We used two X-ray-based thermal pressure profile estimates for MACSJ0647.7+7015: (1) including the over-pressure region (PXO) and (2) excluding the over-pressure region (PX). The latter was already presented in Ruppin et al. (2018). These two pressure profiles allow us to assess the consistency between the X-ray and tSZ views of the ICM. We show in the top row of Fig. 2 the electronic density and temperature as directly derived from the *XMM-Newton* data set. The dark and light green dots (and error bars) correspond to the profiles derived from the full cluster data and excluding the over-pressure, respectively. In the bottom row we show the derived pressure profile with the same color scheme. It should be noted that in this analysis we are limited by the extension of the temperature profile, which can be computed up to about 1.3 Mpc (full data) and 1 Mpc (excluding over-pressure). In the following, the hydrostatic mass profile for the X-ray-only results are extrapolated linearly beyond this radius.

2.3. Estimation of the hydrostatic mass profile

Assuming HE, the total mass enclosed within the radius r is given by

$$M_{\text{HSE}}(r) = - \frac{r^2}{\mu_{\text{gas}} m_p n_e(r) G} \frac{dP_e(r)}{dr}, \quad (2)$$

where G is the gravitational constant, m_p the proton mass, $\mu_{\text{gas}} = 0.61$ the mean molecular weight of the gas, and n_e and P_e the cluster gas electron density and pressure radial profiles.

The hydrostatic mass profile of MACSJ0647.7+7015 computed using Eq. (2) is shown in Fig. 3 for the five pressure profiles discussed above: PN, PNO, PA10, PX, and PXO. In terms of electron density, we used the electron density profiles discussed above including (in combination with PNO, PA10, and PXO) or excluding (in combination with PN and PX) the over-pressure region. As discussed above, the extension of the hydrostatic mass profiles for the X-ray-only estimate was limited by the temperature profiles, and they were then extrapolated linearly beyond 1.3 Mpc and 1 Mpc for the analysis, including and excluding the over-pressure region, respectively. To derive the combined X-ray and tSZ hydrostatic mass profiles, the non-parametric pressure estimates were previously fitted to a gNFW model (see Fig. 1), and estimates of the pressure derivative were obtained from the best-fit models.

We observe significant differences between the three tSZ- and X-ray-based estimates (for PN, PNO, and PA10), even at the large radii at which the cluster masses are traditionally given in the literature, R_{500} and R_{200} . We remind the reader that when including the over-pressure region the tSZ data are significantly not circularly symmetric, thus breaking the hypothesis made for the computation of the pressure and hydrostatic mass profiles. In addition, we find that the two X-ray-only-based estimates (for PXO and PX) are consistent with each other and with the PN-based estimate. The over-pressure does not seem to affect the measured X-ray brightness.

We present in Table 1 estimates of the characteristic radius, R_{500}^{HE} , and of the hydrostatic mass at this radius, M_{500}^{HE} , for the PN, PNO, and PA10 pressure profiles. We notice that the estimates for the PXO and PX pressure profiles are consistent with those of the PN one. We observe in the table that there are significant differences in the estimate of the characteristic mass and the radius, which are either related to the over-pressure region or to the differences in the modeling of the pressure profile. These mass estimates can also be compared to results in the literature based on integrated quantities as $M_{500}^{\text{MCXC}} = 8.8 \times 10^{14} M_\odot$ based on L_{500} (Piffaretti et al. 2011) and $M_{500}^{\text{Planck}} = (8.2 \pm 0.7) \times 10^{14} M_\odot$ based on Y_{SZ} (Planck Collaboration XXVII 2016). Although here we only present results for a single cluster, these differences in the mass estimates are illustrative of the systematic uncertainties one would encounter when deriving the hydrostatic mass of a cluster for which we do not have high resolution observations. These uncertainties need to be accounted for when deriving scaling relations relating the cluster integrated observables (Y_{500} and Y_{500}^{X}) to the hydrostatic mass. In the following, we define R_{500}^{tSZX} as the characteristic radius, R_{500}^{HE} , for the PN pressure profile estimate and use it for comparison with the lensing estimate discussed in Sect. 3.

3. Lensing mass density profile

In this section we discuss the reconstruction of the lensing mass density profile of MACSJ0647.7+7015 using data from the *Hubble* Space Telescope (Postman et al. 2012).

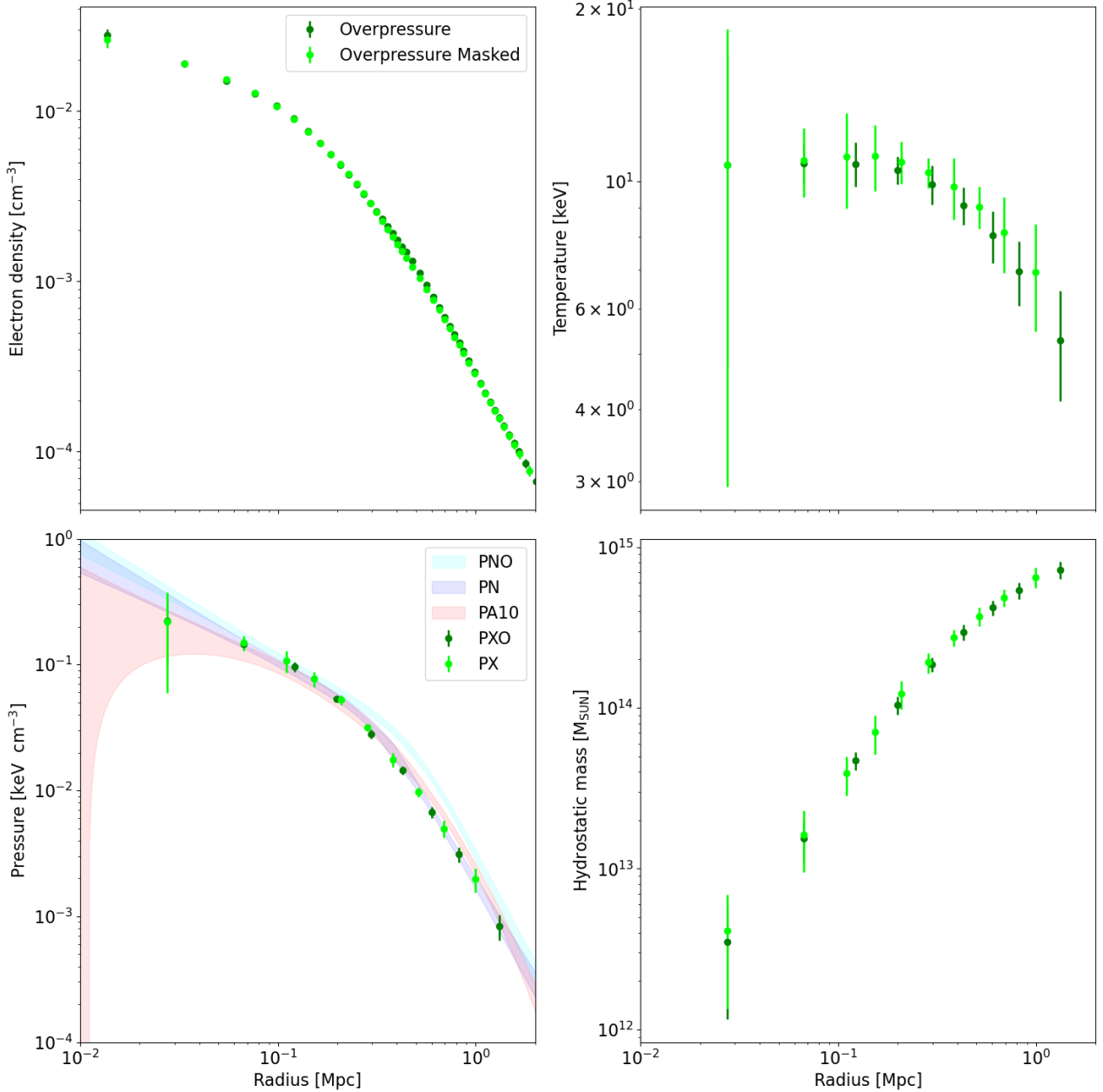


Fig. 2. Electron density, temperature, pressure, and hydrostatic mass profiles as derived from the *XMM-Newton* X-ray data. The dark and light green dots represent the profile computed using the full cluster emission and masking the over-pressure area, respectively. We also show for comparison the PN, PNO, and PA10 pressure profiles presented in Fig. 1.

3.1. Data and preprocessing

Our analysis is based on the joint weak and strong lensing convergence map (hereafter κ -map) of MACS J0647.7+7015 obtained by Zitrin et al. (2015) in their study of the 25 CLASH clusters (Postman et al. 2012). In Zitrin et al. (2015) the authors considered two different parametrizations for the lens model: (1) a light-trace-mass (LTM) density profile and (2) a pseudo isothermal elliptical mass distribution plus an elliptical Navarro-Frenk-White (PIEMD+eNFW) density profile (Navarro et al. 1996). More details on these two approaches can be found in Zitrin et al. (2009, 2013a,b, 2015). Although we considered both parametrizations, we limit our presentation here to the analysis of the κ -map corresponding to the second approach as we find equivalent results. The PIEMD+eNFW κ -map of MACS J0647.7+7015 delivered by Zitrin et al. (2015) exhibits

an elliptical morphology, with a somewhat larger ellipticity than the tSZ NIKA2 map of the cluster, and with the major axis pointing slightly northward.

The convergence, κ , is a measure of the projected mass density per unit of critical density at sky position θ :

$$\kappa(\theta) = \Sigma(\theta)/\Sigma_{\text{cr}} \quad (3)$$

with

$$\Sigma_{\text{cr}} = \frac{c^2}{4\pi G} \frac{D_s}{D_1 D_{1s}}, \quad (4)$$

where D_1 , D_s , and D_{1s} are the angular diameter distance between, respectively, the observer and the lens, the observer and the source, and the source and the lens. The θ is the 2D angular separation with respect to the cluster center.

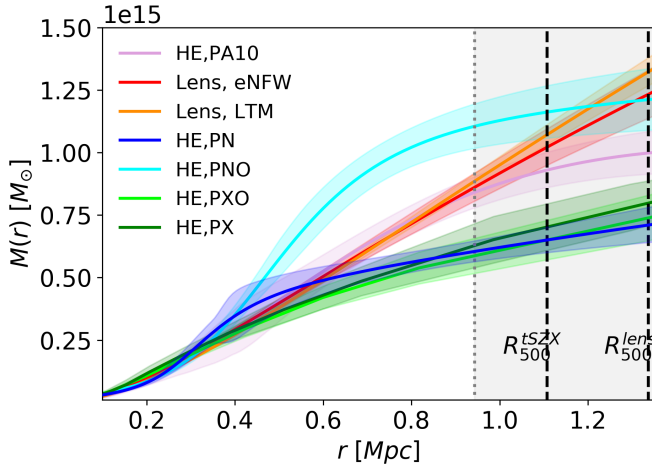


Fig. 3. Combined X-ray and tSZ hydrostatic mass profile computed using three different electron pressure estimates: (1) non-parametric fit using the full 150 GHz NIKA2 map (HE over-pressure, PNO, cyan), (2) non-parametric fit masking the over-pressure region (HE, PN, blue), and (3) the universal [Arnaud et al. \(2010\)](#) pressure profile for this cluster (pink, PA10). Also shown are reconstructed X-ray-only hydrostatic mass profiles, excluding (PX, dark green) and including over-pressure (PXO, bright green), as well as a reconstructed lensing mass profile for the eNFW (red) and LTM models (orange). The vertical dotted line represents the maximum radius at which the lensing data are available. The dashed lines show the characteristic radii R_{500}^{tSZX} and R_{500}^{lens} (see text for details).

Table 1. R_{500} and M_{500} as derived from the hydrostatic mass profiles for the PN, PNO, PA10, PXO, and PX pressure profile estimates.

Pressure profile	R_{500}^{HE} [kpc]	M_{500}^{HE} [$10^{14} M_{\odot}$]
PN	1107 ± 30	6.95 ± 0.56
PNO	1342 ± 61	12.42 ± 1.43
PA10	1236 ± 178	9.7 ± 0.8
PXO	1070 ± 40	6.4 ± 0.8
PX	1110 ± 50	7.0 ± 0.9

Using Eqs. (3) and (4), and assuming that the background sources are at a redshift of $z_s = 2$ ([Zitrin et al. 2011, 2015](#)), we can directly compute the projected mass density profile, $\Sigma(r)$, from the κ -map by averaging the signal in rings of increasing distance from the center of the cluster, which is taken to be the same as the center used in the hydrostatic mass analysis. In the left panel of Fig. 4 we show the projected mass density profile obtained by averaging the PIEMD+eNFW κ -map within 60 rings logarithmically spaced up to a distance from the center of the cluster of 2.3 arcmin, corresponding to a physical radius of $R \sim 1$ Mpc (black points). The uncertainties in the projected mass density profile are computed from the standard deviation in each bin in order to account, to first order, for the observed ellipticity in the PIEMD+eNFW κ -map.

3.2. Three-dimensional mass profile reconstruction

From the projected lensing mass density we can reconstruct the 3D density profile of the cluster, which we parametrize by using a 3D Navarro-Frenk-White (NFW; [Navarro et al. 1996](#)) density profile. Assuming spherical symmetry for the cluster, and using Eqs. (7)–(9) from [Bartelmann \(1996\)](#), the projected 3D NFW

density profile is given by

$$\Sigma(x) = \begin{cases} \frac{2\rho_s r_s}{x^2-1} \left(1 - \frac{2}{\sqrt{x^2-1}} \arctan \sqrt{\frac{x-1}{x+1}} \right) & (x > 1) \\ \frac{2\rho_s r_s}{x^2-1} \left(1 - \frac{2}{\sqrt{1-x^2}} \operatorname{arctanh} \sqrt{\frac{1-x}{1+x}} \right) & (x < 1) \\ \frac{2\rho_s r_s}{3} & (x = 1) \end{cases} \quad (5)$$

as a function of the NFW model parameters ρ_s and r_s , where $x = r/r_s$ and $\rho_s = \delta_c \rho_c$. We define

$$\delta_c = \frac{200}{3} \frac{c_{200}^3}{\ln(1+c_{200}) - c_{200}/(1+c_{200})}, \quad (6)$$

where $c_{200} = r_{200}/r_s$ is the concentration parameter and ρ_c the critical density of the Universe.

We fitted the numerical projected mass density profile from Fig. 4 to the model in Eq. (5) using as free parameters the characteristic cluster radius, r_s , and the concentration parameter, c_{200} . We performed an MCMC analysis to obtain the best-fit parameters. In particular, we used the emcee MCMC software implemented in Python by [Foreman-Mackey et al. \(2013\)](#) and [Goodman & Weare \(2010\)](#). We ran the MCMC code using 100 walkers until the convergence criteria proposed by [Gelman & Rubin \(1992\)](#) were fulfilled for all the model parameters. The 2D and 1D posterior probability distributions for r_s and c_{200} are shown in the right panel of Fig. 4. The marginalized best-fit values are $r_s = 0.56 \pm 0.07$ Mpc and $c_{200} = 3.68 \pm 0.29$. For comparison, we also show in the left panel of Fig. 4 our best-fit model, including 1σ uncertainties.

Using these best-fit parameters and uncertainties, we computed the lensing mass profile of MACS J0647.7+7015, which we display in Fig. 3. For illustration, the corresponding $R_{500}^{\text{lens}} = 1340 \pm 52$ kpc is shown. We also find $R_{200}^{\text{lens}} = 2060 \pm 235$ kpc. From the mass profile we obtain $M_{500}^{\text{PIEMD+eNFW,lens}} = (12.3 \pm 1.4) \times 10^{14} M_{\odot}$ and the virial mass $M_{200}^{\text{PIEMD+eNFW,lens}} = (18.1 \pm 2.4) \times 10^{14} M_{\odot}$. We notice that from the analysis of the LTM κ -map we find slightly larger masses, but they are consistent within uncertainties: $M_{500}^{\text{lens,LMT}} = (13.7 \pm 0.9) \times 10^{14} M_{\odot}$. Combining the two estimates we obtain $M_{500}^{\text{lens}} = (13.0 \pm 2.8) \times 10^{14} M_{\odot}$. We assumed that the difference between the two models is a good representation of the modeling's systematic uncertainties. The derived lensing masses are consistent with the one obtained by [Umetsu et al. \(2014\)](#), $M_{200}^{\text{lens}} = (13.2 \pm 4.2) \times 10^{14} M_{\odot}$; in their analysis they combined the strong and weak lensing constraints of a fraction of 20 CLASH clusters and fitted a stacked density profile.

However, our best-fit value of the concentration parameter is more than 2σ larger than theirs. Similar conclusions can be drawn from the comparison to the results recently obtained by [Umetsu et al. \(2018\)](#), who found $M_{200}^{\text{lens}} = (11.73 \pm 3.79) \times 10^{14} M_{\odot}$ using an eNFW model.

We observe that there is good consistency between the different lensing mass estimates. However, we must stress here the fact that numerical simulations indicate that lensing mass estimates may be affected by systematic uncertainties (see [Pratt et al. 2019](#), for a review). Among them, the uncertainties induced by the mass modeling are probably well represented by the scatter within the different estimates presented here. On the contrary, others, such as those induced by the uncertainties on the redshift and shape of the cluster galaxies and on the choice of background galaxies, are difficult to estimate for the data used in this paper.

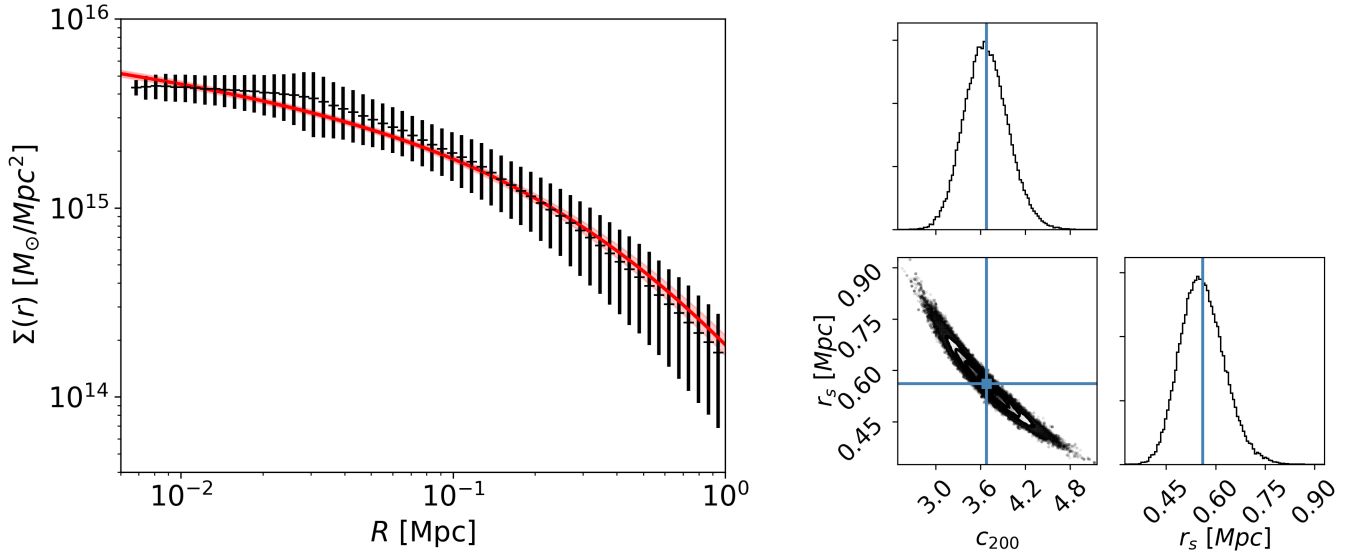


Fig. 4. Lensing analysis. *Left:* projected lensing mass density profile as a function of radius from the PIEMD+eNFW κ -map Zitrin et al. (2015) lens model (black points). Also shown are our best-fit projected NFW profile and 1σ uncertainties (red line and shaded area). The 2D sky angular separation from the cluster center has been converted to radial distance (in Mpc) using the angular distance at the cluster redshift. *Right:* 2D and 1D posterior probability distributions for the c_{200} and r_s parameters of the 3D mass profile as obtained from the MCMC analysis.

4. Discussion

4.1. Gas to lensing mass fraction

From the electron density profile described in Sect. 2.2 we can compute the gas mass profile. Using the lensing mass profile, we computed the gas to lensing mass fraction $f_{\text{gas,lens}} = \frac{M_{\text{gas}}}{M_{\text{lens}}}$, which is presented in Fig. 5 as a function of the normalized radius, assuming R_{500}^{lens} as a reference. As neither the electron density nor the lensing mass profile estimates reach large radii, we believe that the steep increase in the gas fraction observed at those radii is most probably due to a computational artifact. Uncertainties at the 1σ level are shown. We find that at R_{500}^{tSZX} the gas fraction is $f_{\text{gas}} = 0.084 \pm 0.009$. This value is compatible with the results of Pratt et al. (2016) from an X-ray analysis. It is also marginally in agreement with the results found from the MUSIC hydro-dynamical simulations (Sembolini et al. 2013), the results obtained by Chiu et al. (2018) from their analyses of 91 South Pole Telescope selected clusters, and the results of Planck Collaboration Int. V (2013) on intermediate and low redshift clusters.

Furthermore, by assuming a stellar fraction of $\sim 6\%$ (Chiu et al. 2018) and that the lensing mass estimate is a good proxy for the total mass of the cluster, the resulting baryon fraction is in agreement with the one from Planck Collaboration VI (2020).

4.2. Comparison of the hydrostatic and lensing mass profiles

In Fig. 3 we directly compare the five estimates of the hydrostatic mass profiles and the lensing mass profiles derived in Sects. 2.3 and 3. We can observe significant differences between these profiles. We see that in the case of the joint tSZ and X-ray estimates, the over-pressure region significantly increases the hydrostatic mass profile estimate (HE, PNO), as discussed in Ruppin et al. (2018). Such an increase results in a larger cluster hydrostatic mass for a given fixed characteristic cluster radius. The estimate of M_{500} from the lensing analysis is smaller but is consistent with the one from the hydrostatic analysis when including the over-

pressure (HE, PNO), and nearly a factor of two larger than the one obtained when excluding the over-pressure (HE,PN). They are also a factor of two larger than the ones obtained from the X-ray-only estimates, both excluding (HE, PX) and including (HE, PXO) the over-pressure region.

In Fig. 3 we also show the cluster characteristic radius as computed from the hydrostatic mass profile excluding the over-pressure region (HE, PN), R_{500}^{tSZX} , and the one derived from the lensing mass profile, R_{500}^{lens} . We observe that the two estimates are significantly different.

The effects discussed above add an extra layer of complexity to the cosmological use of the hydrostatic mass estimates. They show that in practice: (1) the estimates of R_{500} and, as a consequence, M_{500} may vary significantly, and (2) the details of the dynamical state of the cluster significantly modify the estimates of the cluster properties and their inter-comparison between different cluster observables. The consequences of the presence of the over-pressure are very different on the hydrostatic mass estimates derived from the tSZ and X-ray joint analysis and from the X-ray-only analysis. The study of these systematic effects cannot rely on a single cluster and must be undertaken on a representative sample of clusters. This is one of the major scientific goals of the NIKA2 LPSZ (Mayet et al. 2020; Macias-Pérez et al. 2017).

4.3. Hydrostatic to lensing mass bias

From the results on the gas to lensing mass fraction, it seems that the lensing mass might be a good estimate of the total mass of the cluster. Therefore, it is interesting to infer a hydrostatic to lensing mass bias. In Fig. 6 we show the ratio between the hydrostatic mass profile and the lensing mass profile for the five pressure profile estimates discussed in Sect. 2. We observe that the $M_{\text{HE}}/M_{\text{lens}}$ ratio varies from the inner part of the cluster to the outskirts. In particular, we find a clear peak for the three hydrostatic mass profile estimates based on the joint tSZ and X-ray analysis, which is probably due to the sharp change in slope that biases toward larger values of the hydrostatic mass in the inner part of the cluster. Significant gradients are also observed in the

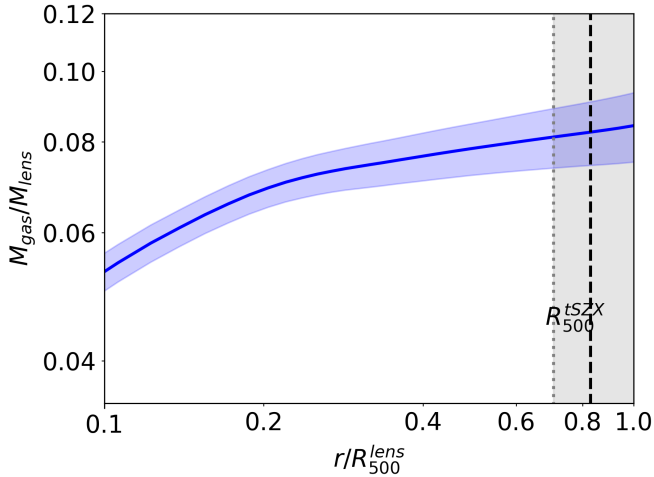


Fig. 5. Gas to lensing mass fraction and uncertainties (blue line and violet shaded area).

$M_{\text{HE}}/M_{\text{lens}}$ ratio between R_{500}^{tSZX} and R_{500}^{lens} , but it becomes constant at very large radii, typically on the order of R_{200} . We note that for these large radii the lensing mass profile is not directly measured but is extrapolated from the best-fit model.

We computed the hydrostatic to lensing mass bias from the $M_{\text{HE}}/M_{\text{lens}}$ ratio for a given characteristic radius, R_{500} . We defined it as $B_{\text{HE-lens}} = \frac{M_{500}^{\text{HE}}}{M_{500}^{\text{lens}}} = (1 - b_{\text{HE-lens}})$ following the notation in [Planck Collaboration XX \(2014\)](#), [Planck Collaboration XXII \(2016\)](#) for the hydrostatic mass bias. Building on the discussion presented in the previous section, we expect this value to vary significantly for the different estimates of the hydrostatic mass. As the estimates of R_{500} vary significantly, and to further illustrate the possible systematic uncertainties in cosmological analyses, we computed this hydrostatic to lensing mass bias for the two estimates of the characteristic radius, R_{500} , discussed above: R_{500}^{tSZX} and R_{500}^{lens} . The main results are given in Table 2. We observe in the table that the hydrostatic to lensing mass bias varies significantly, undergoing both positive and negative bias. These differences are mainly due to the differences in the hydrostatic mass profiles discussed above. However, we also stress the fact that to compute the hydrostatic to lensing mass bias, a common choice of the characteristic radius has to be made, and this can introduce further systematic uncertainties. In Table 2 we observe that for this cluster the hydrostatic to lensing mass bias is consistent within uncertainties for the two extreme choices of the characteristic radius, R_{500} .

5. Summary and conclusions

We have presented in this paper a high resolution multiwavelength analysis of the cluster of galaxies MACSJ0647.7+7015. This analysis is a pilot study for the exploitation of the cluster sample from the NIKA2 LPSZ, which will consist of a sample of 45 clusters (24 of which have been already observed) with redshift ranging from 0.5 to 0.9 and expected masses, M_{500} , from 10^{14} to $10^{15} M_{\odot}$.

We have used high and low resolution tSZ observations from the NIKA2 ([Ruppin et al. 2018](#)) and *Planck* ([Planck Collaboration XXII 2016](#)) experiments, X-ray data from the *XMM-Newton* satellite, and the lensing reconstruction obtained from the *Hubble* Space Telescope CLASH project ([Zitrin et al. 2015](#)). From the projected observables and using

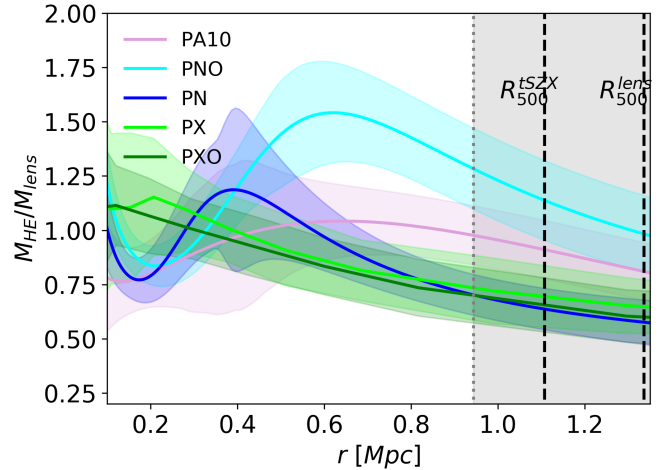


Fig. 6. Hydrostatic to lensing mass ratio for the hydrostatic mass profile estimates for the different pressure profile estimates presented in Sect. 2.3. Uncertainties at the 1σ level are shown as shaded areas of the corresponding color.

Table 2. Hydrostatic to lensing mass bias, $B_{\text{HE-lens}}$, computed from the ratio of the hydrostatic and lensing mass profiles.

HE mass estimates	Hydrostatic to lensing mass bias at	
	R_{500}^{lens}	R_{500}^{tSZX}
PNO	0.98 ± 0.18	1.14 ± 0.18
PN	0.58 ± 0.01	0.64 ± 0.10
PA10	0.81 ± 0.13	0.91 ± 0.13
PXO	0.60 ± 0.12	0.66 ± 0.12
PX	0.65 ± 0.12	0.69 ± 0.14

Notes. We compute it for the five hydrostatic mass profile estimates as well as for $R_{500}^{\text{lens}} = 1342 \pm 52$ kpc and $R_{500}^{\text{tSZX}} = 1107 \pm 30$ kpc.

MCMC techniques, we have been able to derive to high accuracy the 3D hydrostatic and lensing mass profiles beyond the R_{500} characteristic radius of the cluster. Following the analysis of [Ruppin et al. \(2018\)](#), we computed the hydrostatic mass profile including and excluding the over-pressure region that they identified for different combinations of the tSZ and X-ray data and for the universal profile model ([Arnaud et al. 2010](#)). We find significant differences between the different estimates of the hydrostatic mass. These differences are related to the assumptions on the dynamical state of the cluster (including or excluding the over-pressure region), the observables used (tSZ and X-ray joint analysis versus X-ray-only analysis), and the pressure profile models considered.

From the comparison of the hydrostatic and lensing mass profiles of MACSJ0647.7+7015 we discuss systematic uncertainties in the estimates of the characteristic quantities that represent the cluster, such as the characteristic radius (R_{500}), the hydrostatic and lensing masses (M_{500}^{HE} and M_{500}^{lens}) at that radius, their ratio, and the gas fraction. We find that the hydrostatic mass estimates, M_{500} , at the characteristic radius, R_{500} , vary by factors of up to two depending on the dynamical state of the cluster (including or excluding the over-pressure region), the choice of the data used (combined tSZ and X-ray versus X-ray-only analyses), and the pressure profile estimate. We also find that the hydrostatic mass estimates are, as expected, smaller overall than those obtained from the lensing mass profile but for the joint tSZ and X-ray analysis when the over-pressure is considered.

Equivalently, the estimates of R_{500} vary by up to 30%. This makes the comparison between different mass estimates given in the literature difficult if no high resolution mass profiles are available. Moreover, these variations can lead to important systematics in the scaling relations used for cosmological analyses and should be accounted for.

We also computed the MACSJ0647.7+7015 gas to lensing fraction at a fixed characteristic radius, R_{500} , and find that it is in a marginal agreement with the results of [Planck Collaboration Int. V \(2013\)](#) on intermediate and low redshift clusters. However, by assuming that the lensing mass is a good proxy for the total mass of the cluster and assuming a stellar fraction of $\sim 6\%$, the resulting baryon fraction is compatible with the cosmic one ([Planck Collaboration VI 2020](#)).

Finally, we conclude that the above results demonstrate the need for detailed multiwavelength high resolution observations of a representative cluster sample at high redshift, as expected from the NIKA2 LPSZ, in order to fully characterize the mass-observable scaling relations used in cosmological cluster analyses.

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