



HAL
open science

The spin-orbit alignment of the transiting exoplanet WASP-3b from Rossiter-McLaughlin observations

E. K. Simpson, D. Pollacco, G. Hébrard, N. P. Gibson, S. C. C. Barros, I. Boisse, F. Bouchy, A. Collier Cameron, G. R. M. Miller, C. A. Watson, et al.

► **To cite this version:**

E. K. Simpson, D. Pollacco, G. Hébrard, N. P. Gibson, S. C. C. Barros, et al.. The spin-orbit alignment of the transiting exoplanet WASP-3b from Rossiter-McLaughlin observations. *Monthly Notices of the Royal Astronomical Society*, 2010, 405, pp.1867-1872. 10.1111/j.1365-2966.2010.16576.x . hal-03646109

HAL Id: hal-03646109

<https://hal.science/hal-03646109v1>

Submitted on 3 Jun 2022

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

The spin–orbit alignment of the transiting exoplanet WASP-3b from Rossiter–McLaughlin observations[★]

E. K. Simpson,¹ † D. Pollacco,¹ G. Hébrard,² N. P. Gibson,³ S. C. C. Barros,¹ I. Boisse,² F. Bouchy,^{2,4} A. Collier Cameron,⁵ G. R. M. Miller,⁵ C. A. Watson¹ and F. P. Keenan¹

¹*Astrophysics Research Centre, School of Mathematics & Physics, Queen's University, University Road, Belfast BT7 1NN*

²*Institut d'Astrophysique de Paris, UMR7095 CNRS, Université Pierre & Marie Curie, 98 bis bd. Arago, 75014 Paris, France*

³*School of Physics, University of Exeter, Stocker Road, Exeter EX4 4QL*

⁴*Observatoire de Haute-Provence, CNRS/OAMP, 04870 St Michel l'Observatoire, France*

⁵*School of Physics and Astronomy, University of St Andrews, North Haugh, St Andrews, Fife KY16 9SS*

Accepted 2010 February 20. Received 2010 February 19; in original form 2009 December 18

ABSTRACT

We present an observation of the Rossiter–McLaughlin effect for the planetary system WASP-3. Radial velocity measurements were made during transit using the SOPHIE spectrograph at the 1.93-m telescope at Haute-Provence Observatory. The shape of the effect shows that the sky-projected angle between the stellar rotation axis and planetary orbital axis (λ) is small and consistent with zero within 2σ ; $\lambda = 15^{+10}_{-9}^\circ$. WASP-3b joins the \sim two-thirds of planets with measured spin–orbit angles that are well aligned and are thought to have undergone a dynamically gentle migration process such as planet–disc interactions. We find a systematic effect which leads to an anomalously high determination of the projected stellar rotational velocity ($v \sin i = 19.6^{+2.2}_{-2.1} \text{ km s}^{-1}$) compared to the value found from spectroscopic line broadening ($v \sin i = 13.4 \pm 1.5 \text{ km s}^{-1}$). This is thought to be caused by a discrepancy in the assumptions made in the extraction and modelling of the data. Using a model developed by Hirano et al. designed to address this issue, we find $v \sin i$ to be consistent with the value obtained from spectroscopic broadening measurements ($v \sin i = 15.7^{+1.4}_{-1.3} \text{ km s}^{-1}$).

Key words: stars: planetary systems – stars: individual: WASP-3 – techniques: radial velocities – techniques: photometric.

1 INTRODUCTION

Bright transiting planets hold the key to determining properties of exoplanet systems which are otherwise inaccessible. One such example is the projected obliquity of the planet's orbit relative to the rotation axis of the star (λ). This can be measured from spectroscopic observations during transit, which reveal a radial velocity deviation from the standard Keplerian orbital motion known as the Rossiter–McLaughlin (RM) effect (McLaughlin 1924; Rossiter 1924). The effect is caused by the planet passing over and blocking a portion of the rotating stellar surface. This introduces an asymmetry in the spectral line profile and a small apparent shift in the position of the line. The resulting shape of the RM effect traces the trajectory of the planet across the stellar disc and, once modelled, allows the projected spin–orbits alignment to be found.

The alignment angle provides an insight into the history of these planetary systems. Hot Jupiters are thought to have formed far out in the system and have migrated inwards. Migration via planet–disc interactions is believed to be relatively non-violent and would not perturb the orbital obliquity (Lin, Bodenheimer & Richardson 1996; Murray et al. 1998). It is thought that the time-scale for spin–orbit alignment via tidal interactions is of a similar order to the lifetime of the system (Barker & Ogilvie 2009; Levrard, Winisdoerffer & Chabrier 2009) and therefore that the primordial alignment of the system is preserved. Assuming that a planet was formed in an aligned protoplanetary disc, it too is expected to be well aligned.¹ However, more dynamically violent processes such as planet–planet scattering and Kozai oscillations from interactions with a stellar companion are thought to significantly effect the orbital obliquity and could produce highly misaligned and even retrograde orbit (e.g. Rasio & Ford 1996; Wu & Murray 2003; Nagasawa, Ida &

[★]Based on observations collected with the SOPHIE spectrograph on the 1.93-m telescope at Observatoire de Haute-Provence (CNRS), France, by the SOPHIE Consortium (program 08B.PNP.SIMP).

†E-mail: esimpson05@qub.ac.uk

¹Several authors have noted that this may not necessarily be the case in all systems, especially those which have undergone interactions with other stars (e.g. Golinowski et al. 2006; Bate, Lodato & Pringle 2010).

Bessho 2008). Thus the distribution of alignment angles allows us to constrain the formation and evolution mechanisms of planetary systems.

The form of the RM effect depends on λ as well as the projected equatorial rotation velocity of the star ($v \sin i$) and several parameters which can be measured from a transit light curve (see Section 2.2). For the WASP-3 system, the amplitude of the effect can be estimated as $\sim 0.7 v \sin i (R_p/R_*)^2 \sqrt{1-b^2} \sim 86 \text{ m s}^{-1}$, where R_p and R_* are the planetary and stellar radii, b is the impact parameter and 0.7 is a scaling factor which approximates the effects due to limb darkening. WASP-3 is the fastest rotator for which the RM effect has not been previously reported ($v \sin i = 13.4 \text{ km s}^{-1}$; Pollacco et al. 2008). It is also a bright target ($V_{\text{mag}} = 10.6$), and therefore an ideal candidate for measuring the RM effect. WASP-3 is a main-sequence star of spectral type F7-8V. It hosts a hot Jupiter of mass $1.8M_J$ and radius $1.3R_J$ in a close orbit of 1.8 d (Gibson et al. 2008; Pollacco et al. 2008). Because of the close proximity and relative temperature of its host star, WASP-3b is one of the most highly irradiated exoplanets known.

It has been shown that the formulae used to describe the RM effect (Ohta, Taruya & Suto 2005, hereafter OTS) appear to overestimate $v \sin i$ compared to the value found from spectral line broadening, and the deviation scales with $v \sin i$ (e.g. Winn et al. 2005; Triaud et al. 2009). This has led Hirano et al. (2010, hereafter H10) to re-evaluate the equations assuming that the distortion of the radial velocity shift is being measured with a cross-correlation technique (as used in this work). We apply this method to the RM effect of WASP-3 and compare it to the results obtained using the model derived by OTS.

In this paper we report on time-series spectroscopic observations of WASP-3 obtained during transit. Section 2 describes the observations and analysis procedure performed to measure the spin-orbit angle of the system. The results of the derived system parameters are presented in Section 3, and discussed in Section 4.

2 OBSERVATIONS AND METHOD

2.1 Observations and data reduction

A transit of WASP-3b was observed with the SOPHIE spectrograph at the 1.93-m telescope at Haute-Provence Observatory on the night of 2008 September 30. SOPHIE is a cross-dispersed, environmentally stabilized echelle spectrograph (wavelength range 3872.4–6943.5 Å) designed for high-precision radial velocity measurements (Bouchy et al. 2009). The spectrograph was used in high-efficiency mode (resolution $R = 40\,000$) and the CCD in fast read-out mode. Two 3-arcsec diameter optical fibres were used, the first centred on the target and the second on the sky to simultaneously measure its background in case of contamination from scattered moonlight. The moon illumination was 3 per cent and at a distance of $> 86^\circ$ on the night of transit so this did not significantly affect the radial velocity determination.

We acquired 26 spectra of WASP-3 on the night of transit, covering the full transit (137 min) and a period of duration 130 min out of transit. The exposure times range from 5 to 31 min in order to reach a constant signal-to-noise ratio of 35 at 550 nm, reflecting the variable throughputs obtained due to changing atmospheric conditions. The sky was clear when the observations began, and the airmass ranged from 1.1 to 2.7 during the sequence with increasing cloud during the later hours. The exposures remained sufficiently short to provide a good time sampling with 17 measurements taken during transit. In addition, we used seven out-of-transit observations

at various orbital phases obtained during the discovery of the planet in 2007 to fit the orbit (Pollacco et al. 2008).

The spectra were extracted and radial velocities measured using the SOPHIE pipeline (Perruchot et al. 2008; Bouchy et al. 2009). Radial velocities were computed from a weighted cross-correlation of each spectrum with a numerical mask of G2 spectral type, as described by Baranne et al. (1996) and Pepe et al. (2002). A Gaussian was fitted to the cross-correlation functions to obtain the radial velocity shift, full width at half-maximum (FWHM) and contrast with respect to the continuum. The uncertainty in the radial velocity was computed using the empirical relation given in Bouchy et al. (2005) and Cameron et al. (2007). Uncertainties were typically $\sim 38 \text{ m s}^{-1}$ during the RM sequence and $\sim 23 \text{ m s}^{-1}$ at other orbital phases. For the observations taken on the night of transit, the spectrum below 4367 Å was removed from the cross-correlation in an effort to reduce the systematic errors caused by the airmass effect. The wavelength stability over the duration of the observations was $\sim 3 \text{ m s}^{-1}$. The seven orbital points reported in Pollacco et al. (2008) were re-reduced using an updated version of the pipeline. Table 1 summarizes the measured radial velocities.

Table 1. Radial velocities and 1σ error bars of WASP-3 measured with SOPHIE during and outside transit. The orbital points are updated from Pollacco et al. (2008) using a new version of the reduction pipeline.

BJD –240 0000	RV (km s^{-1})	Error (km s^{-1})
Planetary transit		
54740.29509	–5.42487	0.03856
54740.29909	–5.38158	0.03628
54740.30328	–5.36639	0.03751
54740.30771	–5.43318	0.03734
54740.31312	–5.45738	0.03756
54740.31830	–5.40972	0.03828
54740.32351	–5.42139	0.03816
54740.32907	–5.33800	0.03782
54740.33524	–5.35707	0.03764
54740.34152	–5.38229	0.03741
54740.34699	–5.38596	0.03626
54740.35300	–5.36107	0.03533
54740.35891	–5.36891	0.03508
54740.36506	–5.41906	0.03437
54740.37170	–5.49966	0.03580
54740.37771	–5.55305	0.03686
54740.38426	–5.55488	0.03571
54740.39228	–5.58417	0.03825
54740.39993	–5.64022	0.03606
54740.40758	–5.58523	0.03519
54740.41466	–5.69115	0.03415
54740.42259	–5.64015	0.03543
54740.43066	–5.53109	0.03945
54740.44171	–5.59288	0.03844
54740.45894	–5.62403	0.04132
54740.48062	–5.52085	0.05355
Other orbital phases		
54286.52255	–5.72553	0.02137
54287.45639	–5.18430	0.03541
54289.36622	–5.18792	0.03047
54340.32511	–5.64313	0.01280
54341.39898	–5.40394	0.01914
54342.31981	–5.54401	0.02801
54343.48259	–5.60354	0.01283

2.2 Determination of system parameters

The RM effect and orbit were fitted simultaneously using all the available spectroscopic data. A Keplerian model was used for the orbit, and the analytical approach described in OTS for the RM effect. The OTS equations were modified to make them dependent on R_p/R_* and a/R_* rather than R_p , R_* and a , to reflect the parameters derived from photometry, and reduce the number of free parameters. The model comprises 13 parameters: orbital period, P ; mid-transit time, T_0 ; planetary to stellar radius ratio, R_p/R_* ; scaled semimajor axis a/R_* ; orbital inclination, i ; orbital eccentricity, e ; longitude of periastron, ω ; radial velocity semi-amplitude, K ; systemic velocity of orbital data set, γ_1 ; systemic velocity of transit data set, γ_2 ; sky projected angle between the stellar rotation axis and orbital angular momentum vector, λ ; projected stellar rotational velocity, $v \sin i$ and stellar linear limb-darkening coefficient, u .

The six observations taken immediately before the transit ingress indicate that the systemic velocity of the transit data set is $\sim 25 \text{ m s}^{-1}$ lower than that measured from the orbital velocities. We therefore allow the systemic velocity of the orbit and transit data to differ (γ_1 and γ_2). All of the out of transit points (from 2007 and 2008) were first fitted and no significant deviation from a circular orbit was found ($e = 0.07 \pm 0.08$). The eccentricity was then fixed to zero for the rest of the analysis.

We used a linear limb darkening law, as the quadratic law alters the model by less than $\sim 3 \text{ m s}^{-1}$ and so does not seem justified by the precision of the RM data, especially given that for $b \sim 0.5$ as in this case, the alignment angle is primarily determined by the time when the radial velocity anomaly vanishes, rather than the exact form of the model. The linear limb-darkening coefficient $u = 0.69$ was chosen from the tables of Claret (2004) (ATLAS models) for the g' filter. We used the value corresponding to $T_{\text{eff}} = 6500 \text{ K}$, $\log g = 4.5$, $[\text{M}/\text{H}] = 0.0$ and $\zeta = 2 \text{ km s}^{-1}$ (a typical value for this spectral type; see Gray, Graham & Hoyt 2001) which most closely represented the properties of WASP-3 ($T_{\text{eff}} = 6400 \text{ K}$, $\log g = 4.25$, $[\text{M}/\text{H}] = 0.0$, Pollacco et al. 2008). As a test, we set u as free parameter and found that there was no significant effect on λ or $v \sin i$.

We placed no constraint on $v \sin i$ as there is a known systematic effect in its determination from RM measurements (see Section 3.2). Triaud et al. (2009) suggest that there may be a similar systematic effect influencing R_p/R_* if solely spectroscopic observations are used.² We therefore fixed the photometric parameters (P , T_0 , R_p/R_* , a/R_* and i) to those found through light-curve modelling. A Markov chain Monte Carlo (MCMC) fit to a full and partial transit of WASP-3 was presented in Gibson et al. (2008). In this analysis, the stellar radius was calculated based on the scaling relation $R_* \propto M_*^{1/3}$ with a prior on M_* , leading to an underestimation in the error on the orbital inclination. The light curves were refitted with the stellar radius as a free parameter using the procedure described in Gibson et al. (2010). Five separate chains each with two-hundred thousand points were computed. To test that the chains had all converged to the same parameter space, the Gelman & Rubin statistic (Gelman & Rubin 1992) was calculated for each of the free parameters. It was found to be less than 1 per cent from unity for all parameters, a good sign of mixing and convergence. The results are shown in Table 2.

² This is because the amplitude of the RM effect is primarily determined by R_p/R_* and $v \sin i$, so the observed increased amplitude could lead to a false inflation of these values.

Table 2. Adopted system parameters and uncertainties required for the RM effect, and other photometric parameters updated from Gibson et al. (2008).

Parameter (units)	Symbol	Value
Period (d)	P	1.846835 ± 0.000002
Transit epoch (HJD)	T_0	$245\,4605.55915 \pm 0.00023$
Orbital inclination ($^\circ$)	i	$84.93^{+1.32}_{-0.78}$
Planet/star radius ratio	R_p/R_*	$0.1013^{+0.0014}_{-0.0013}$
Scaled semimajor axis	a/R_*	$5.173^{+0.246}_{-0.162}$
Eccentricity ^a	e	0 (adopted)
Limb darkening ^b	u	0.69
Impact parameter	b	$0.466^{+0.047}_{-0.111}$
Transit duration (h)	T_d	$2.737^{+0.024}_{-0.022}$
Stellar radius (R_\odot)	R_*	$1.31^{+0.05}_{-0.07}$
Planet radius (R_J)	R_p	$1.29^{+0.05}_{-0.07}$

^aSee Section 2.

^bFrom Claret (2004).

Thus the model contains five free parameters (λ , $v \sin i$, K , γ_1 and γ_2). The χ^2 statistic was calculated:

$$\chi^2 = \sum_i \left[\frac{v_{i,\text{obs}} - v_{i,\text{calc}}}{\sigma_i} \right]^2, \quad (1)$$

where $v_{i,\text{obs}}$ and $v_{i,\text{calc}}$ are the i th observed and calculated radial velocities and σ_i the corresponding observational error. The optimal parameters were obtained by minimizing the χ^2 statistic using the IDL-based MPFIT function (Markwardt 2009); a least-squares minimization technique using the Levenberg–Marquardt algorithm. The starting values for the parameters were those derived in Pollacco et al. (2008) as well as $\lambda = 0$ and $\gamma_2 = \gamma_1$.

The spectroscopic errors were rescaled so that the best-fitting model has a reduced χ^2 (χ^2_{red}) of 1, which required the initial spectroscopic errors to be multiplied by the factor $\sqrt{\chi^2_{\text{red}}} = 0.905$. This suggests that the errors were slightly overestimated however the rescaling did not significantly affect the best-fitting parameters or errors.

The 1σ best-fitting parameter uncertainties were calculated using a Monte Carlo method. We created 10^5 synthetic data sets by adding a 1σ Gaussian random variable to the data points. The photometric parameters were also varied for each realization by taking a parameter set from the distributions from the light-curve fitting. This allowed correlations between parameters such as a/R_* and i to be taken in to account. The starting values for the free parameters were also perturbed by a 3σ Gaussian random variable to further explore the parameter space. The free parameters were re-optimised for each simulated data set to obtain the distribution of the best-fitting parameter values. The distributions were not assumed to be Gaussian and the 1σ limits were found where the distribution enclosed ± 34.1 per cent of the values away from the median. In order to verify the errors, we also ran a global fit to both the photometric and radial velocities using an MCMC algorithm, as described in Anderson et al. (2010).

3 RESULTS

3.1 System parameters

The derived parameters and errors for the WASP-3 system are shown in Table 3 for both the least-squares and MCMC methods. The results from the global MCMC analysis are consistent with those

Table 3. Best-fitting values and uncertainties using OTS (both least-squares and MCMC fitting) and H10 models.

Parameter (units)	Symbol	OTS	OTS (MCMC)	H10
Projected alignment angle ($^{\circ}$)	λ	15_{-9}^{+10}	14 ± 9	13_{-7}^{+9}
Projected stellar rotation velocity (km s^{-1})	$v \sin i$	$19.6_{-2.1}^{+2.2}$	19.6 ± 2.2	$15.7_{-1.3}^{+1.4}$
RV semi-amplitude (km s^{-1})	K	0.274 ± 0.011	0.273 ± 0.013	0.276 ± 0.011
Systemic velocity of orbit data set (km s^{-1})	γ_1	-5.458 ± 0.008	-5.458 ± 0.005	-5.458 ± 0.007
Systemic velocity of transit data set (km s^{-1})	γ_2	-5.483 ± 0.011	-5.486 ± 0.011	-5.487 ± 0.009
Chi-squared (spectroscopic data)	χ^2	27.9	27.4	27.2

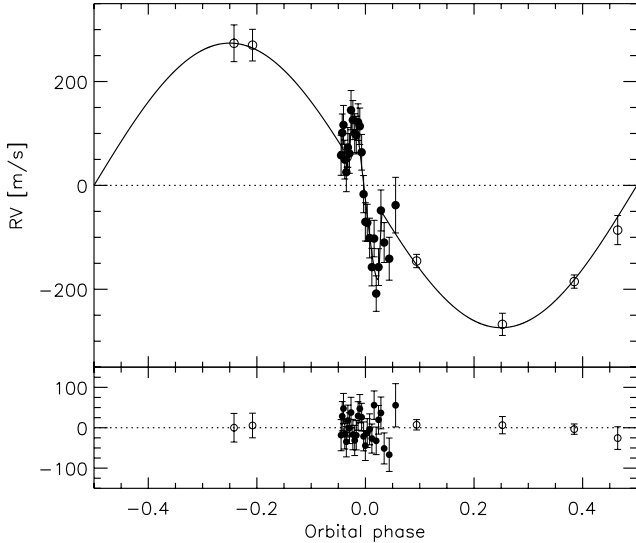


Figure 1. Upper panel: radial velocities of WASP-3 minus the systemic velocity, overplotted with the best-fitting model. Open points are from Pollacco et al. (2008). Filled points show the RM effect of WASP-3 on 2008 September 30. Lower panel: residuals from the best-fitting curve.

found from the least-squares fitting. The re-reduced radial velocities yield orbital parameters which are consistent with Pollacco et al. (2008).

Fig. 1 shows all the WASP-3 radial velocities minus the systemic velocity, the best-fitting model and the residuals to the fit. Fig. 2 shows a zoom-in on the observations taken on the night of transit minus the Keplerian orbit, the best-fitting model and residuals. It can be seen that there is a ~ 10 -min offset between the radial velocity and photometric zero-points (solid and dotted lines). The elapsed time between the transit photometry and the RM observations is approximately 73 orbits and the 0.173 s uncertainty in P and 20 s uncertainty in T_0 causes an uncertainty of only 33 s which cannot explain the offset. By fitting the data, we obtain $\lambda = 15_{-9}^{+10}$, indicating that the timing offset is caused by a small misalignment between the planetary orbit and stellar spin axis. The system is generally well aligned and the misalignment is not highly significant, with λ consistent with zero to within 2σ .

The difference between the offset velocities of the seven orbital data points and those taken on the night of transit is significant to 2σ ; $\sim 28 \pm 14 \text{ m s}^{-1}$, so we investigated modelling all the data together without including a relative shift. We found that the newly derived parameters were consistent with the previous values to within the error bars. The value of λ increased to 23_{-7}^{+8} because λ is highly correlated with γ . The new fit yields a γ velocity which is higher than that found for the transit data set alone, so the RM effect appears to cross the zero velocity point earlier, and produces a larger

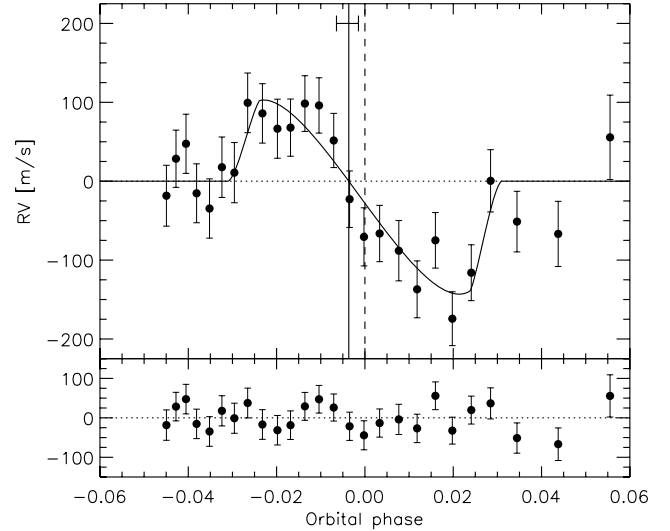


Figure 2. Upper panel: RM effect of WASP-3 minus the Keplerian motion, overplotted with the best-fitting model. The solid and dotted lines show the offset between the radial velocity and photometric zero-points, suggesting that the system is not completely aligned. The error bar on the solid line shows the 1σ uncertainty in the position of the radial velocity zero-point. The uncertainty on the photometric zero-point is smaller than the width of the line. Lower panel: residuals from the best-fitting model.

misalignment angle. There are several factors which could cause an offset between the data sets. These include a long-term trend caused by an additional body, stellar activity, instrumental drift or a combination of these.³ In addition, the data taken on the night of transit underwent a slightly different reduction procedure which may well cause a shift in the radial velocities between nights (see Section 2) so we prefer to fit the two data sets separately.

3.2 Stellar rotational velocity and systematics

The observed amplitude of the RM effect ($\sim 100 \text{ m s}^{-1}$) is significantly larger than expected ($\sim 86 \text{ m s}^{-1}$, see Section 1). Fig. 3 shows the predicted RM effect with $v \sin i$ equal to that found from modelling the spectral line broadening ($v \sin i = 13.4 \pm 1.5 \text{ km s}^{-1}$; Pollacco et al. 2008). By fitting the RM effect, we find the projected rotational velocity to be overestimated by ~ 50 per cent; $v \sin i = 19.6_{-2.1}^{+2.2} \text{ km s}^{-1}$. Several authors have noted a similar discrepancy, which is particularly evident in faster rotating stars. Winn et al. (2005) suggested that systematics are present in the determination of the radial velocities because the asymmetric stellar line profile is being fitted with a symmetric Gaussian. They attempted to correct

³ The shift is larger than the typical value for astrophysical jitter for this spectral type so this is unlikely to be the sole cause.

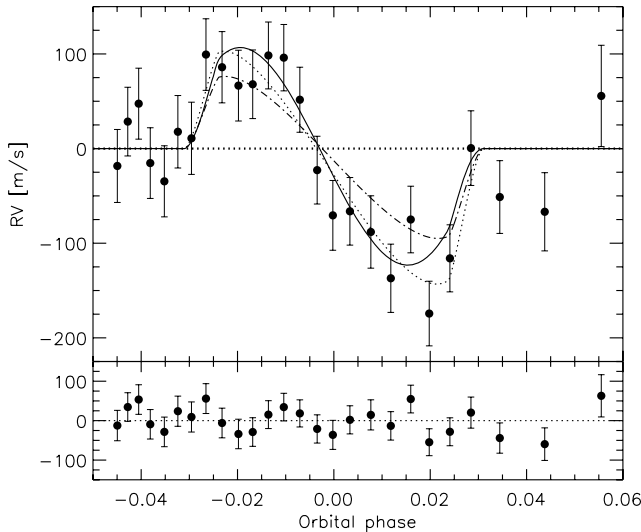


Figure 3. Upper panel: RM effect of WASP-3 with various models. Solid (red) is the fit using the H10 model with $v \sin i$ as a free parameter ($v \sin i = 15.3 \text{ km s}^{-1}$). Dotted (blue) is the OTS model with $v \sin i$ as a free parameter ($v \sin i = 19.3 \text{ km s}^{-1}$). Dash-dot (blue) is the fit using the OTS model but with $v \sin i$ held fixed at $v \sin i = 13.4 \text{ km s}^{-1}$ as determined through spectroscopic broadening (Pollacco et al. 2008). The amplitude of this model is obviously too low. All the models cross the x-axis at the same phase, giving consistent values of λ . Lower panel: residuals from the H10 best-fitting model.

this by adding a quadratic term to the RM model which significantly improved the estimated $v \sin i$.

Triaud et al. (2009) showed that in HD 189733, the structure of the residuals observed when subtracting the RM model from the data was consistent with those found when subtracting the RM model from simulated data undergoing the same data reduction procedure. This indicates that the cause of the discrepancy is indeed in the extraction of the radial velocities. The residuals of the WASP-3 RM effect show a very similar structure during the first three quarters of the transit to those in fig. 4 of Triaud et al. (2009) (the last section is dominated by systematics caused by airmass and weather effects). The amplitude of the residuals is of the order of the error bars so it is not clear whether the structure is detected significantly.

The apparent radial velocity shift of the stellar lines predicted by the OTS model is produced by the distortion of the line profiles by the planet blocking of a velocity component of the star's spectrum. Assuming that the resulting asymmetry of the stellar lines is not well resolved, a symmetric function fitted to the line profile measures a shift in the mean position. For faster rotating stars, the asymmetry in the line profile is better resolved because of the broader width, and causes a larger apparent shift than expected, as seen here.

H10 have addressed this problem by modifying the OTS equations under the assumption that the RV shift is measured by cross-correlation with a Gaussian function. They find that the RM model is also dependent on the intrinsic width of the individual line profile.

Hence, instead of modelling the RM velocity anomaly using the OTS equation:

$$\delta v = f v_p / (1 - f), \quad (2)$$

where f is flux blocked by planet $\sim (R_p/R_*)^2$ and v_p is subplanet velocity (i.e. the velocity component of the rotating stellar surface blocked by the planet $\sim x_p v \sin i/R_*$, where x_p is the x coordinate of the position of the planet on the stellar surface, see fig 5 of OTS),

we now use the H10 model (equations 41 and 42):

$$\delta v = f v_p (p - q v_p^2) = f v_p (1.72 - 0.00546 v_p^2), \quad (3)$$

where

$$p = \left[1 + \frac{\sigma^2}{2\beta^2 + \sigma^2} \right]^{3/2}, \quad q = \frac{p}{2\beta^2 + \sigma^2} \quad (4)$$

and β is the intrinsic linewidth, which has been calculated by Miller et al. (in preparation) for WASP-3 as $\beta = 11.1 \pm 0.1 \text{ km s}^{-1}$. The quantity $\sigma = v \sin i/\alpha$, where α is a scaling factor depending on limb darkening parameters (see equation F6 of H10). We calculate $\alpha = 1.31$ for the fixed limb darkening coefficients $u_1 = 0.69$ and $u_2 = 0$.

We fit this equation to the data and find that λ is consistent with that found using the OTS model, $\lambda = 13_{-7}^{+9^\circ}$ ($\lambda = 15_{-9}^{+10^\circ}$; OTS). This is expected because λ mainly depends on the shape of the waveform and the time crossing zero-point rather than the amplitude. The slight difference is probably caused by fit finding a marginally lower value of γ_2 , which is correlated with λ . We find $v \sin i$ to be much closer to the line-broadening value $v \sin i = 15.3 \pm 1.4 \text{ km s}^{-1}$. It is still slightly different but is acceptable considering the errors. There was no significant difference in χ^2 between the OTS and H10 fits. Fig. 3 shows the data overplotted with the best-fitting OTS and H10 models.

We also investigated whether the discrepancy between the two independent measurements of $v \sin i$ could be caused by differential rotation. Gaudi & Winn (2007) show that if the surface speed of a star is dependant on latitude as well as longitude, then the form of the RM effect is slightly altered and depends on the fractional difference in rotation speed between the pole and equator (α). To reproduce the $\sim 6 \text{ km s}^{-1}$ difference in $v \sin i$ seen in WASP-3, α must be greater than 0.45. Reiners (2006) showed that no stars have been observed with $\alpha > 0.45$ so this scenario seems unlikely. In addition, the form for the differential rotation law would have to be antisolar, which is seldom observed. The bias caused by the fitting of Gaussians to asymmetric CCFs offers a much more viable explanation.

4 CONCLUSIONS

A full transit of WASP-3b was observed with the SOPHIE spectrograph at the 1.93-m telescope at Haute-Provence Observatory. We modelled the RM effect using the equations of OTS and found the sky projected spin-orbit alignment angle of the system to be $\lambda = 15_{-9}^{+10^\circ}$. WASP-3b is not strongly misaligned and is consistent with 0° to within 2σ . This suggests that WASP-3b underwent a relatively non-violent migration process which did not perturb it from the primordial alignment of the protoplanetary disc. The result indicates a marginal detection of a small misalignment, and further, higher precision observations of this target would establish the reality of this.

We find a significant discrepancy between the $v \sin i$ value from the RM model ($v \sin i = 19.6_{-2.1}^{+2.2} \text{ km s}^{-1}$) and the measurement from spectroscopic line broadening analysis ($v \sin i = 13.4 \pm 1.5 \text{ km s}^{-1}$; Pollacco et al. 2008). This phenomenon has been noted in several systems, especially for fast rotators. It is thought to be the result of systematics introduced when measuring a line profile asymmetry as a shift in the mean line position.

H10 provide a modified formulation to take into account the method used to determine radial velocity shifts. We applied this model and found the alignment angle to be consistent with the OTS model, and derived an improved value for the rotational velocity

$v \sin i = 15.7_{-1.3}^{+1.4} \text{ km s}^{-1}$. This is consistent with the value found from spectroscopic line broadening analysis. The RM effect has the potential to provide a much more precise determination of $v \sin i$ than spectral broadening for slow rotators, as it is a geometrical measurement that is not subject to uncertainties such as macroturbulence. However, RM models must be shown to ensure that the $v \sin i$ discrepancy is fully resolved.

Since the beginning of 2009, the number of systems reported with tilts significantly greater than 30° has risen dramatically from one (XO3; Hebrard et al. 2008) to six (CoRoT-1, Pont et al. 2010; HAT-P-7, Narita et al. 2009; Winn et al. 2009b; HD 80606, Gillon 2009; Moutou et al. 2009; Pont et al. 2009; Winn et al. 2009a; WASP-14, Johnson et al. 2009; WASP-17, Anderson et al. 2010). The RM effect has now been measured for a total of 20 systems and WASP-3 joins the $\sim 2/3$ of these which are well aligned. These statistics are re-enforcing the inference by Fabrycky & Winn (2009) that there is a bimodal distribution of spin-orbit angles, with a fraction of systems being well aligned while the rest have random mutual orientations, suggesting that there are multiple migration mechanisms at work. It is therefore vital to increase the number of measured systems, both aligned and misaligned, to place these theories on a secure statistical basis.

ACKNOWLEDGMENTS

We thank the technical team at Haute-Provence Observatory for their support with the SOPHIE instrument and the 1.93-m OHP telescope. FPK is grateful to AWE Aldermaston for the award of a William Penny Fellowship. EKS would like to thank T. Hirano for his help with the modified RM equations and R. M. Crockett for his support and nursing of injuries during the observing campaign.

REFERENCES

- Anderson D. R. et al., 2010, *ApJ*, 709, 159
 Baranne A. et al., 1996, *A&AS*, 119, 373
 Barker A. J., Ogilvie G. I., 2009, *MNRAS*, 395, 2268
 Bate M. R., Lodato G., Pringle J. E., 2010, *MNRAS*, 401, 1505
 Bouchy F., Pont F., Melo C., Santos N. C., Mayor M., Queloz D., Udry S., 2005, *A&A*, 431, 1105
 Bouchy F. et al., 2009, *A&A*, 505, 853
 Cameron A. C. et al., 2007, *MNRAS*, 375, 951
 Claret A., 2004, *A&A*, 428, 1001
 Fabrycky D. C., Winn J. N., 2009, *ApJ*, 696, 1230
 Gaudi B. S., Winn J. N., 2007, *ApJ*, 655, 550
 Gelman A., Rubin D., 1992, *Statistical Sci.*, 7, 457
 Gibson N. P. et al., 2008, *A&A*, 492, 603
 Gibson N. P. et al., 2010, *MNRAS*, 401, 1917
 Gillon M., 2009, preprint (arXiv:0906.4904)
 Golimowski D. A. et al., 2006, *AJ*, 131, 3109
 Gray R. O., Graham P. W., Hoyt S. R., 2001, *AJ*, 121, 2159
 Hebrard G. et al., 2008, *A&A*, 488, 763
 Hirano T., Suto Y., Taruya A., Narita N., Sato B., Johnson J. A., Winn J. N., 2010, *ApJ*, 709, 458 (H10)
 Johnson J. A., Winn J. N., Albrecht S., Howard A. W., Marcy G. W., Gazak J. Z., 2009, *PASP*, 121, 1104
 Lestrade B., Winisdoerffer C., Chabrier G., 2009, *ApJ*, 692, L9
 Lin D. N. C., Bodenheimer P., Richardson D. C., 1996, *Nat*, 380, 606
 McLaughlin D. B., 1924, *ApJ*, 60, 22
 Markwardt C. B., 2009, in Bohlender D. A., Durand D., Dowler P., eds, *ASP Conf. Ser. Vol. 411, Astronomical Data Analysis Software and Systems XVIII*. Astron. Soc. Pac., San Francisco, p. 251
 Moutou C., Hébrard G., Bouchy F., Eggenberger A., Boisse I., 2009, *A&A*, 498, L5
 Murray N., Hansen B., Holman M., Tremaine S., 1998, *Sci*, 279, 69
 Nagasawa M., Ida S., Bessho T., 2008, *ApJ*, 678, 498
 Narita N., Sato B., Hirano T., Tamura M., 2009, *Publ. Astron. Soc. Jpn. Lett.*, 61, 35
 Ohta Y., Taruya A., Suto Y., 2005, *ApJ*, 622, 1118 (OTS)
 Pepe F., Mayor M., Galland F., Naef D., Queloz D., Santos N. C., Udry S., Burnet M., 2002, *A&A*, 388, 632
 Perruchot S., Kohler D., Bouchy F., Richaud Y., Richaud P. e., 2008, *Proc. SPIE*, 7014, 17
 Pollacco D. et al., 2008, *MNRAS*, 385, 1576
 Pont F. et al., 2009, *A&A*, 502, 695
 Pont F. et al., 2010, *MNRAS*, 402, 1
 Rasio F. A., Ford E. B., 1996, *Sci*, 274, 954
 Reiners A., 2006, *A&A*, 446, 267
 Rossiter R. A., 1924, *ApJ*, 60, 15
 Triard A. H. M. J. et al., 2009, *A&A*, 506, 377
 Winn J. N. et al., 2005, *ApJ*, 631, 1215
 Winn J. N. et al., 2009a, *ApJ*, 703, 2091
 Winn J. N., Johnson J. A., Albrecht S., Howard A. W., Marcy G. W., 2009b, *ApJ*, 703, L99
 Wu Y., Murray N., 2003, *ApJ*, 589, 605

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.