



HAL
open science

A translucent interstellar cloud at $z = 2.69$. CO, H₂, and HD in the line-of-sight to SDSS J123714.60+064759.5

P. Noterdaeme, P. Petitjean, C. Ledoux, S. López, R. Srianand, S. D. Vergani

► To cite this version:

P. Noterdaeme, P. Petitjean, C. Ledoux, S. López, R. Srianand, et al.. A translucent interstellar cloud at $z = 2.69$. CO, H₂, and HD in the line-of-sight to SDSS J123714.60+064759.5. *Astronomy and Astrophysics - A&A*, 2010, 523, 10.1051/0004-6361/201015147 . hal-03646027

HAL Id: hal-03646027

<https://hal.science/hal-03646027v1>

Submitted on 4 Sep 2022

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

A translucent interstellar cloud at $z = 2.69$

CO, H₂, and HD in the line-of-sight to SDSS J123714.60+064759.5^{*}

P. Noterdaeme¹, P. Petitjean², C. Ledoux³, S. López¹, R. Srianand⁴, and S. D. Vergani^{5,6}

¹ Departamento de Astronomía, Universidad de Chile, Casilla 36-D, Santiago, Chile
e-mail: [pasquier;slopez]@das.uchile.cl

² Université Paris 6, Institut d'Astrophysique de Paris, CNRS UMR 7095, 98bis bd Arago, 75014 Paris, France
e-mail: petitjean@iap.fr

³ European Southern Observatory, Alonso de Córdova 3107, Vitacura, Casilla 19001, Santiago 19, Chile
e-mail: cledoux@eso.org

⁴ Inter-University Centre for Astronomy and Astrophysics, Post Bag 4, Ganeshkhind, 411 007 Pune, India
e-mail: anand@iucaa.ernet.in

⁵ Université Paris 7, APC, CNRS UMR 7164, 10 rue Alice Domon et Léonie Duquet, 75205 Paris Cedex 13, France

⁶ GEPI, Observatoire de Paris, CNRS UMR 8111, 5 place Jules Janssen, 92195 Meudon, France
e-mail: vergani@apc.univ-paris7.fr

Received 3 June 2010 / Accepted 2 August 2010

ABSTRACT

We present the analysis of a sub-damped Lyman- α system with neutral hydrogen column density, $\log N(\text{H}^0) (\text{cm}^{-2}) = 20.0 \pm 0.15$ at $z_{\text{abs}} = 2.69$ toward SDSS J123714.60+064759.5 ($z_{\text{em}} = 2.78$). Using the VLT/UVES and X-shooter spectrographs, we detect H₂, HD, and CO molecules in absorption with $\log N(\text{H}_2, \text{HD}, \text{CO}) (\text{cm}^{-2}) = 19.21^{+0.13}_{-0.12}$, 14.48 ± 0.05 and 14.17 ± 0.09 respectively. The overall metallicity of the system is super-solar ($[\text{Zn}/\text{H}] = +0.34$ relative to solar) and iron is highly depleted ($[\text{Fe}/\text{Zn}] = -1.39$), revealing metal-rich and dusty gas. Three H₂ velocity components spanning $\sim 125 \text{ km s}^{-1}$ are detected. The strongest H₂ component, at $z_{\text{abs}} = 2.68955$, with $\log N(\text{H}_2) = 19.20$, does *not* coincide with the centre of the HI absorption. This implies that the molecular fraction in this component, $f_{\text{H}_2} = 2N(\text{H}_2)/(2N(\text{H}_2)+N(\text{H}^0))$, is higher than the mean molecular fraction $\langle f_{\text{H}_2} \rangle = 1/4$ in the system. We also found the Cl⁰ associated with this H₂ component to have $N(\text{Cl}^0)/N(\text{Cl}^+) > 0.4$, which points in the same direction. Cl⁰ is tied to H₂ by charge exchange reactions, this means that the molecular fraction in this component is not far from unity. The kinetic temperature derived from the $J = 0$ and 1 rotational levels of H₂ is $T = 108^{+84}_{-33}$ K and the particle density derived from the C⁰ ground-state fine structure level populations is $n_{\text{H}0} \sim 50\text{--}60 \text{ cm}^{-3}$. We derive an electronic density $< 2 \text{ cm}^{-3}$ for a UV field similar to the Galactic one and show that the carbon-to-sulphur ratio in the cloud is close to the solar ratio. The size of the molecular cloud is probably smaller than 1 pc. Both the CO/H₂ = 10^{-5} and CO/C⁰ ~ 1 ratios for $f_{\text{H}_2} > 0.24$ indicate that the *cloud* classifies as translucent, i.e., a regime where carbon is found both in atomic and molecular form. The corresponding extinction, $A_V = 0.14$, albeit lower than the definition of a translucent *sightline* (based on extinction properties), is high for the observed H⁰ column density. This means that intervening clouds with similar local properties but with higher column densities (i.e. larger physical extent) could be missed by current magnitude-limited QSO surveys. The excitation of CO is dominated by radiative interaction with the cosmic microwave background radiation (CMBR) and we derive $T_{\text{ex}}(\text{CO}) = 10.5^{+0.8}_{-0.6}$ K when $T_{\text{CMBR}}(z = 2.69) = 10.05$ K is expected. We measure $N(\text{HD})/2N(\text{H}_2) = 10^{-5}$. This is about 10 times higher than what is measured in the Galactic ISM for $f_{\text{H}_2} = 1/4$ but similar to what is measured in the Galactic ISM for higher molecular fractions. The astration factor of deuterium with respect to the primordial D/H ratio is only about 3. This can be the consequence of accretion of unprocessed gas from the intergalactic medium onto the associated galaxy. In the future, it will be possible to search efficiently for molecular-rich DLAs/sub-DLAs with X-shooter, but detailed studies of the physical state of the gas will still need UVES observations.

Key words. cosmology: observations – galaxies: ISM – quasars: absorption lines – quasars: individual: SDSS J123714.60+064759.5

1. Introduction

Studies of the interstellar medium (ISM) in the local Universe have shown that the neutral ISM presents a complex structure, with cold and dense clouds immersed in a warmer and more diffuse medium. These different ISM phases should be detectable at high redshift by their absorption signatures in damped Lyman- α (DLA) systems observed in quasar spectra (Petitjean et al. 1992).

^{*} Based on observations carried out with X-shooter and the Ultraviolet and Visual Echelle Spectrograph (UVES), both mounted on the European Southern Observatory Very Large Telescope Unit 2 – Kueyen, under Program IDs 082.A-0544(A), 083.A-0454(A) and 084.A-0699(A).

However, although there is evidence of the multiphase nature of DLA systems (e.g. Wolfe et al. 2004), most of the intervening DLAs probe only warm ($T \gtrsim 3000$ K) and diffuse ($n_{\text{H}} < 1 \text{ cm}^{-2}$) atomic gas (e.g. Petitjean et al. 2000; Kanekar & Chengalur 2003). The reason is that the cross-sections of the different phases are quite different and it is not possible to sample them equally well.

Searching for molecular hydrogen in high-redshift DLAs (Ledoux et al. 2003; Petitjean et al. 2006; Noterdaeme et al. 2008a) is an efficient way of detecting colder and denser neutral gas and to probe its physical conditions (e.g. Reimers et al. 2003; Cui et al. 2005; Hirashita & Ferrara 2005; Srianand et al. 2005; Ledoux et al. 2006b; Noterdaeme et al. 2007a,b). Past

Table 1. Journal of observations.

Instrument	Date	Setting	Exposure time	Resolving power ^a	SNR ^b
UVES	2009-03-27	390+564	2 × 5400 s	390B:51400 – 564L:50800 – 564U:49500 – 775L:50200 – 775U:48600	10–40
UVES	2009-03-29	390+564	2 × 5400 s		
UVES	2009-04-27	390+775	2 × 4500 s		
X-shooter	2010-02-24	UVB	3600 s	5100	50
X-shooter	2010-03-03	UVB	3600 s		

Notes. ^(a) “B”, “L” and “U” stand for respectively blue, lower red and upper red CCD. ^(b) Per pixel.

studies have shown that molecular hydrogen is confined in small clouds (pc-sized) with densities $n \sim 1\text{--}100 \text{ cm}^{-3}$ and temperatures $T \sim 70\text{--}200 \text{ K}$. The filling factor of H_2 -bearing clouds in DLAs is much less than one, and only 10% of the lines of sight through a DLA galaxy do intercept H_2 -bearing clouds down to a limit of $N(\text{H}_2) \sim 10^{14} \text{ cm}^{-2}$ (Noterdaeme et al. 2008a). H_2 -bearing clouds in DLAs have small physical extents. Direct evidence for this is that the intervening H_2 -bearing gas does not completely fill the beam from the broad line region of the quasar Q 1232+082 (Ivanchik et al. 2010; Balashev et al. 2010). Nonetheless, the molecular fraction in DLAs remains small and typical of what is seen in Galactic diffuse atomic gas with $f_{\text{H}_2} = 2N(\text{H}_2)/(2N(\text{H}_2) + N(\text{H}^0)) < 0.1$ and often much lower than this (Ledoux et al. 2003; Noterdaeme et al. 2008a).

Snow & McCall (2006) have classified Galactic interstellar clouds into the following categories: (i) *diffuse atomic*, with low molecular fractions; (ii) *diffuse molecular*, where the fraction of hydrogen in molecules becomes substantial ($f_{\text{H}_2} > 0.1$) but carbon is still mainly in ionised form (C^+); (iii) *translucent* (first introduced by van Dishoeck & Black 1989), where the carbon makes the transition to molecular; and (iv) *dense molecular*, where both hydrogen and carbon are fully molecular. As discussed above, most of the H_2 -bearing DLAs detected so far are part of the first, and maybe for some of them, part of the second categories. The fourth category may be difficult to detect in absorption because of the high extinction such a cloud produces on the background source.

Despite their highly interesting chemistry and their close connection with star formation, we know very little about translucent clouds (i.e., the third category) at high redshift. The small cross-section of these clouds and/or the induced extinction of the light from the background sources can probably explain the absence of detection in more than three decades of QSO absorption-line research. However, observing molecular-rich gas in absorption should be possible by selecting sightlines passing through or starting from star-forming regions.

Since long-duration gamma-ray bursts (GRBs) are known to occur within star-forming regions, absorption lines at the host-galaxy redshifts which are imprinted in GRB optical afterglow spectra (e.g., Fynbo et al. 2009) are obvious targets towards this goal. Nevertheless, current samples are characterised by a general lack of H_2 detection (e.g., Fynbo et al. 2006; Tumlinson et al. 2007). This is probably due to the still limited sample sizes as well as to a bias against dusty – molecular-rich – lines of sight (Ledoux et al. 2009; Fynbo et al. 2009). The first detection of both H_2 and CO in the low-resolution spectrum of a highly reddened GRB afterglow (Prochaska et al. 2009) seems to confirm this scenario. Moreover, the observed molecular excitation is high in this case, indicating strong UV pumping from the GRB afterglow itself.

With the large number of quasar spectra available in the Sloan Digital Sky Survey (SDSS), it becomes possible to select

the rare sightlines passing through *intervening* molecular-rich gas. However, due to the small cross-section of such clouds, an efficient selection must be applied. In the local ISM, carbon is found to transition from an ionised state (C^+) to neutral (C^0) and molecular form (CO) from the most superficial to the deepest parts of the clouds (e.g. Snow & McCall 2006; Burgh et al. 2010). From our Very Large Telescope (VLT) survey for H_2 in DLAs (Ledoux et al. 2003; Noterdaeme et al. 2008a), it appears that C^0 is generally observed in the same components as H_2 . This is because the photo-ionisation potential of C^0 is similar to the energy of photons that dissociate H_2 . However, the neutral fraction of carbon is generally low, probably because the gas is not completely shielded. This explains the non-detection of CO in these H_2 -bearing DLAs, even down to $N(\text{CO}) \sim 10^{12} \text{ cm}^{-2}$ (e.g. Petitjean et al. 2002). Searching for systems with high column densities of neutral carbon could be an efficient way to select more shielded gas where other molecules can survive, without relying on a pre-selection based on the H^0 column density (i.e. the absorbers need not be DLAs). Because several C I lines are located redwards of the Lyman- α forest, it is possible to search for strong C I absorptions directly in SDSS spectra with automatic procedures. We therefore initiated a programme to survey specifically these sightlines with the VLT. Our selection has been very successful and already allowed us to detect carbon monoxide along QSO sightlines for the first (Srianand et al. 2008b) and second times (Noterdaeme et al. 2009a). We present here the third detection of CO, at $z = 2.69$ towards SDSS J123714.60+064759.5 ($z_{\text{em}} = 2.78$, hereafter called J 1237+0647). This is a beautiful and peculiar case, for which a detailed analysis of the physical properties of the gas is possible. We present the observations in Sect. 2, the measurement and results in Sect. 3 and provide some discussion in Sect. 4. We summarise our findings in Sect. 5.

2. Observations

2.1. X-shooter observations

We are conducting an observing campaign with X-shooter mounted on the Cassegrain focus of the VLT Unit 2-Kueyen telescope to study the molecular content of our complete sample of C^0 absorbers. As a test case for the sensitivity of X-shooter in the blue we observed J 1237+0647 ($g = 19.2$) twice in service mode on February 24 (airmass 1.2; seeing 1.4'') and March 3, 2010 (airmass 1.3; seeing 1.2''), using a slit width of 1'' in the UVB arm. Each observation run consisted in 1 h exposure taken in staring mode (see Table 1). This yields the nominal resolution power of $R = 5100$ in the UVB arm and a signal-to-noise ratio of about 50 at $\sim 500 \text{ nm}$. Data were reduced with version 0.9.5 of the preliminary ESO X-shooter pipeline (Goldoni et al. 2006) and the appropriate calibration data. The two individual spectra were then combined weighting each pixel by the inverse of the error variance. A portion of the X-shooter spectrum

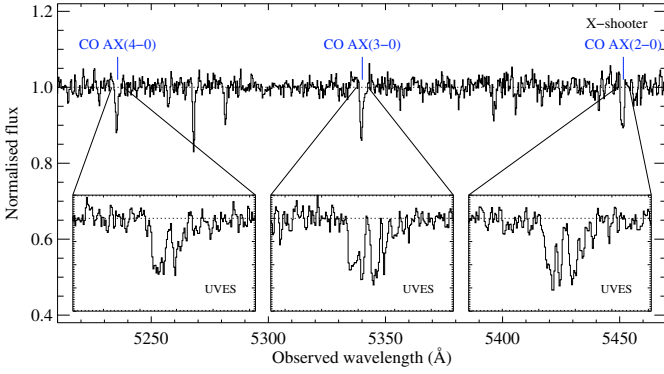


Fig. 1. Portion of X-shooter spectrum in the UVB. The inset figures show 5 Å-wide portions of the UVES spectrum around the position of the detected electronic bands of CO.

featuring CO is shown in Fig. 1, where several electronic bands of CO are clearly detected. These bands are resolved into individual rotational levels in the UVES spectrum obtained with 8.5 h of exposure time (inset figures). From this, it is apparent that X-shooter is the most efficient instrument to survey a complete sample of candidates down to quasar magnitudes as faint as $r \sim 21.5$, whereas using UVES would be excessively time consuming. Then, but only in the case of detection, higher spectral resolution is needed to make a detailed analysis of the physical state of the gas as done in the following.

2.2. UVES observations

The quasar J1237+0647 was observed in visitor mode on March 27 and 29, 2009 and April 27, 2009 with the Ultraviolet and Visual Echelle Spectrograph (UVES, Dekker et al. 2000), mounted at the Nasmyth B focus of VLT-UT2. The total exposure time on-source is 8.5 h (see Table 1). We used two dichroic settings (4×5400 s with 390+564 and 2×4500 s with 390+775) to cover the wavelength range 3300–9600 Å with small gaps at 4517–4621, 5597–5677 and 7764–7809 Å. The CCD pixels were binned 2×2 and the slit width adjusted to $1''$, yielding a resolving power of $\sim 50\,000$ under seeing conditions of 0.9–1''. Individual science spectra were reduced with the ESO UVES pipeline, which performs accurate sky subtraction while removing cosmic ray impacts at the same time. The spectra were then combined with a dedicated IDL routine by weighting each pixel by the inverse of the error variance in that pixel and clipping residual cosmic ray impacts that remained after the cleaning of 2D spectra.

3. Analysis

The system at $z = 2.69$ towards J1237+0647 features numerous absorption lines from atomic (H^0 , O^0 , C^0 , Mg^0 , Cl^0 and S^0), singly-ionised (Fe^+ , Si^+ , Zn^+ , Ni^+ , S^+ , C^+), and molecular species (two isotopomers of molecular hydrogen: H_2 and HD; as well as carbon monoxide: CO).

We analysed the UVES spectrum using standard Voigt profile fitting techniques. The fits were performed through χ^2 -minimisation with the code FITLYMAN (Fontana & Ballester 1995) which is available as a context of the ESO-MIDAS data analysis software. The spectrum was normalised in the wavelength ranges of interest by fitting spline functions to regions free from absorption lines. Atomic data were taken from Morton (2003) for metal lines, unless otherwise specified. Wavelengths

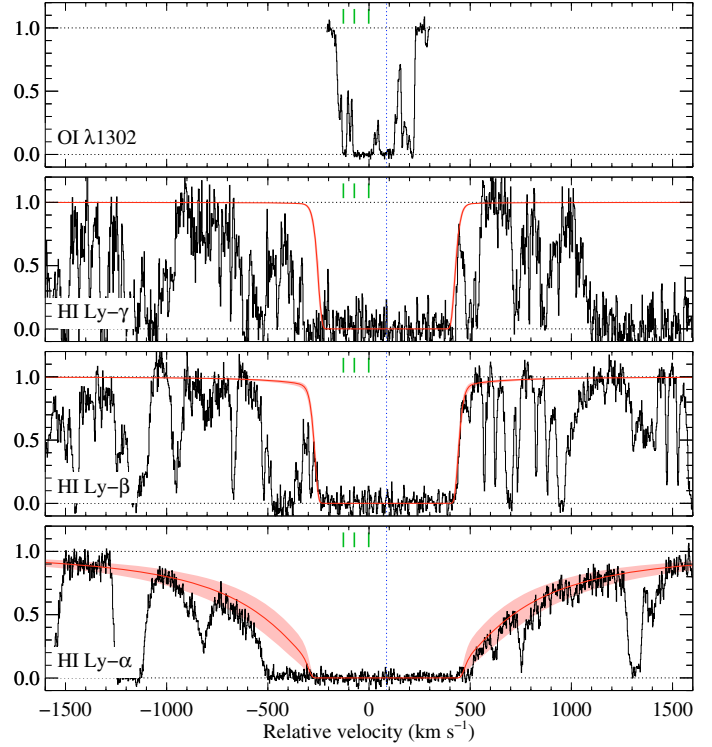


Fig. 2. Measurement of the total column density of neutral atomic hydrogen at $z = 2.69$ towards J1237+0647. The solid line represents the best one-component fit to the data. The centroid of the profile is indicated by a dotted vertical line. Uncertainties on the column density measurement are illustrated by the shaded area. The short tick marks indicate the position of the three H_2 -bearing components, the reddest of which also features CO and HD absorptions. The origin of the velocity scale for this figure and all following ones is defined at the position of the CO-bearing component at $z_{\text{abs}} = 2.68957$.

and oscillator strengths were taken from Morton & Noreau (1994) and Eidelsberg & Rostas (2003) for CO and from Abgrall & Roueff (2006) for HD. Updated wavelengths of H_2 Lyman and Werner bands were taken from Bailly et al. (2010), with oscillator strengths from the Meudon group¹, based on calculations described in Abgrall et al. (1994). Photospheric solar abundances are taken from Asplund et al. (2009).

3.1. Atomic hydrogen

From the damped Lyman- α absorption line (see Fig. 2) we measured the total column density of atomic hydrogen to be $\log N(\text{H}^0) (\text{cm}^{-2}) = 20.00 \pm 0.15$, which agrees with the value measured automatically by Noterdaeme et al. (2009b) from the low resolution SDSS spectrum (20.15 ± 0.28). The centroid of the H I profile is well constrained by the Lyman- β and Lyman- γ absorption lines. The large Doppler parameter ($b \sim 100 \text{ km s}^{-1}$) required to fit the Lyman- β and Lyman- γ lines is a consequence of the presence of multiple components as testified by the clumpy profile of the O I $\lambda 1302$ absorption line spread over $\sim 350 \text{ km s}^{-1}$ (see Fig. 2). We recall that O^0 closely follows H^0 because of favourable charge-exchange reaction. Unfortunately, because of strong saturation and blending effects, it is impossible to derive column densities in individual components and only the total H^0 column density along the line of sight is accessible.

¹ <http://amrel.obspm.fr/molat/>

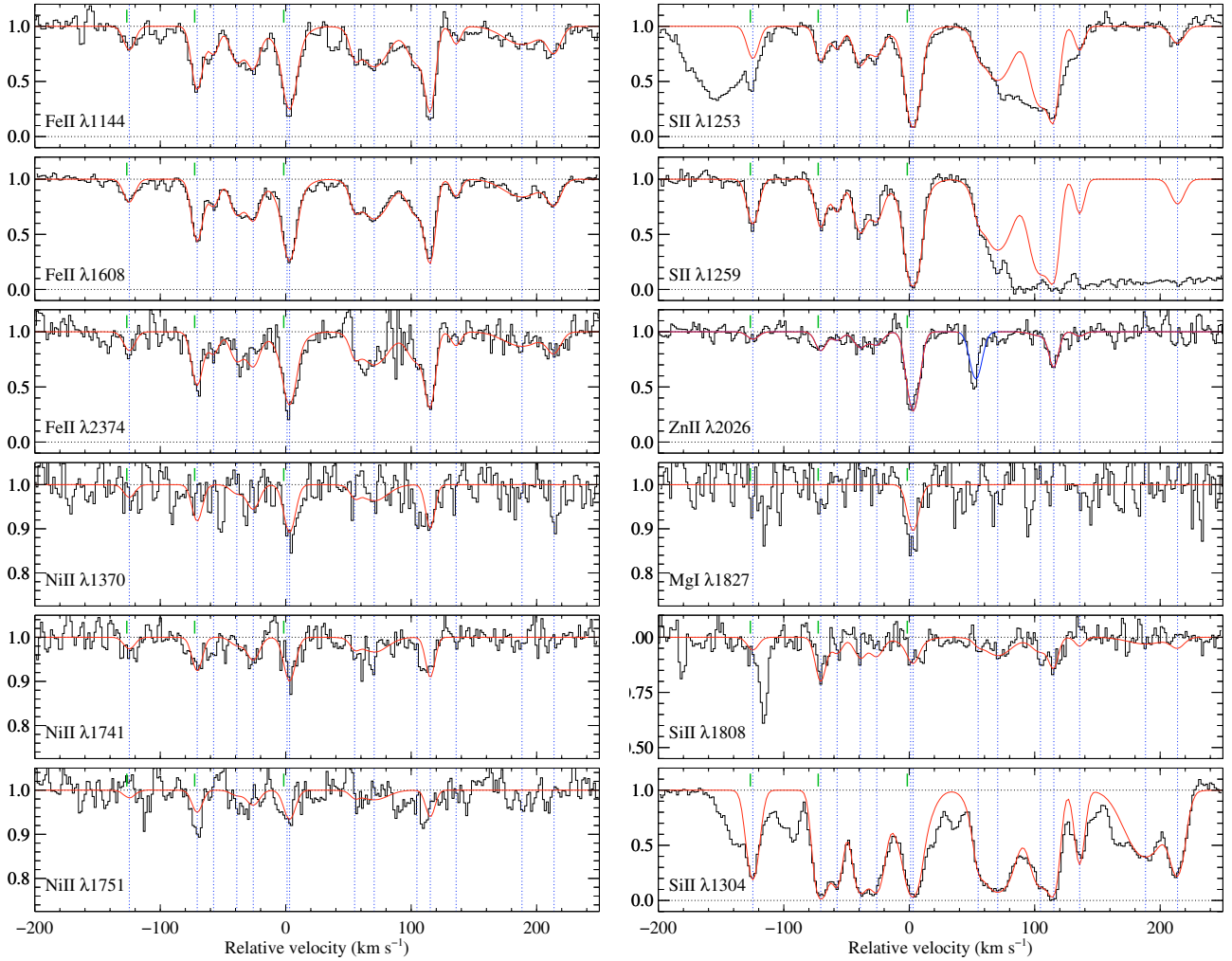


Fig. 3. Fit to metal lines. The origin of the velocity scale is set at the redshift of the CO absorption ($z = 2.68957$). The positions of the three H_2 components are indicated by short tick marks. The absorption seen at $v = +55 \text{ km s}^{-1}$ in the $\text{Zn II } \lambda 2026$ panel (in blue) is due to $\text{Mg I } \lambda 2026$ while that at $v = -110 \text{ km s}^{-1}$ on the $\text{Si II } \lambda 1808$ panel is due to $\text{Si I } \lambda 1807$. Note that both Si II lines are affected by blends. Finally $\text{Si II } \lambda 1304$ reveals the presence of additional weak components that are not fitted (see text).

In Fig. 2 and subsequent figures and tables, the zero of the velocity scale is taken at the position of the CO component ($z_{\text{abs}} = 2.68957$, see Sect. 3.5) and the centroid of the three detected H_2 components at $z_{\text{abs}} = 2.68801, 2.68868$ and 2.68955 ($\Delta v = -127, -73, -1.6 \text{ km s}^{-1}$, see Sect. 3.3) are indicated by short vertical marks. Interestingly, the centroid of the atomic hydrogen absorption profile (vertical dotted line in Fig. 2 at $z_{\text{abs}} = 2.69063$) is shifted by about $+86 \pm 10 \text{ km s}^{-1}$ relative to the CO absorption feature. This, the clumpy O I profile, and the high value of the b -parameter of the H I lines all indicate that a significant fraction of the atomic gas is *not* associated with the molecular gas. We will discuss this below in more detail.

3.2. Metal content

3.2.1. Singly ionised species

Absorption lines from detected low-ionisation species are spread over about 350 km s^{-1} around the strongest component, which is also the component where CO and HD are detected. We used non-saturated transitions to derive accurate column densities for Fe^+ , Ni^+ , S^+ , and Zn^+ . Lines from other species (C^+ , O^0 , N^0) are heavily saturated, which prevented us to derive any meaningful

value of the corresponding column densities. It is however possible to perform an accurate measurement of the Si^+ column density from the simultaneous use of $\text{Si II } \lambda 1808$, which is very weak (below the 3σ detection level), and $\text{Si II } \lambda 1304$, which is close to saturation. Note that although the strong $\text{Si II } \lambda 1304$ line reveals additional weak components, their contribution to the overall column density is negligible. The result of the Voigt-profile fitting is shown in Fig. 3 with the corresponding parameters in Table 2. Both the $\text{Si II } \lambda 1253$ and $\text{Si II } \lambda 1259$ profiles are partially blended. Unfortunately, $\text{Si II } \lambda 1250$ falls in a gap of the spectrum. Therefore, we provide only upper-limits for the blended components in Table 2. $3\text{-}\sigma$ upper-limits on the column densities of Ni^+ and Zn^+ for the undetected components are provided in the table. Finally, we measure $\log N(\text{Cr}^+) < 12.2$ at 3σ from the non-detection of $\text{Cr II } \lambda 2056$.

Total column densities and corresponding mean metallicities in the gas relative to solar are given in Table 3. Note that we do not apply ionisation correction because the presence of neutral and molecular species in the strongest components indicates that the effect of ionisation on the overall abundances should be negligible. Indeed, even in the general population of absorbers, the ionisation correction is only about 0.1 dex for $N(\text{H}^0) = 10^{20} \text{ cm}^{-2}$ (Péroux et al. 2007). The metallicity is

Table 2. Column densities of metal species.

z_{abs}	Δv^a (km s $^{-1}$)	b (km s $^{-1}$)	$\log N^b$ (cm $^{-2}$)					
			Fe $^+$	Ni $^+$	S $^+$	Zn $^+$	Si $^+$	Mg 0
2.68803	-125	5.2 ± 0.2	13.03 ± 0.06	12.18 ± 0.13	14.03 ± 0.02	11.46 ± 0.15	13.82 ± 0.02	<12.70
2.68870	-71	4.6 ± 0.1	13.56 ± 0.04	12.62 ± 0.05	14.04 ± 0.02	11.86 ± 0.06	14.37 ± 0.04	<12.70
2.68886	-57	4.8 ± 0.3	13.09 ± 0.05	<12.15	13.83 ± 0.02	11.50 ± 0.13	≤ 13.91	<12.70
2.68909	-39	6.2 ± 0.5	13.30 ± 0.04	12.07 ± 0.19	14.18 ± 0.01	11.86 ± 0.06	≤ 14.08	<12.70
2.68925	-26	6.0 ± 0.3	13.37 ± 0.04	12.52 ± 0.07	14.01 ± 0.02	11.76 ± 0.08	14.04 ± 0.09	<12.70
2.68958	+1	14.5 ± 1.5	13.41 ± 0.05	<12.15	14.17 ± 0.07	<11.55	13.94 ± 0.04	<12.85
2.68961	+3	5.4 ± 0.2	13.74 ± 0.02	12.79 ± 0.04	14.91 ± 0.02	12.75 ± 0.02	14.04 ± 0.09	13.03 ± 0.02
2.69024	+55	3.5 ± 0.3	12.90 ± 0.06	<12.15	≤ 13.50	<11.40	13.21 ± 0.08	<12.70
2.69044	+71	14.8 ± 2.0	13.69 ± 0.02	12.67 ± 0.09	≤ 14.70	<11.55	14.37 ± 0.01	<12.85
2.69086	+105	9.3 ± 2.0	13.46 ± 0.03	<12.20	≤ 14.80	11.40 ± 0.32	14.09 ± 0.02	<12.75
2.69099	+115	3.7 ± 0.2	13.75 ± 0.02	12.65 ± 0.04	≤ 14.70	12.12 ± 0.04	14.08 ± 0.14	<12.70
2.69124	+136	3.4 ± 0.5	12.81 ± 0.04	<12.15	≤ 13.80	11.15 ± 0.29	13.58 ± 0.04	<12.70
2.69189	+188	19.5 ± 1.7	13.40 ± 0.03	<12.40	<13.40	<11.60	14.04 ± 0.05	<12.95
2.69220	+214	6.5 ± 0.5	13.12 ± 0.03	<12.15	13.78 ± 0.04	<11.40	13.78 ± 0.02	<12.70

Notes. ^(a) In all figures and tables, the velocity is given with respect to the redshift of the CO component at $z = 2.68957$. ^(b) Upper-limits due to blends or saturated lines are indicated by “ \leq ”, while 3σ upper-limits from non-detections are indicated by “ $<$ ”.

Table 3. Summary of overall gas-phase abundances.

Species	$\log N$ (cm $^{-2}$)	Mean abundance ^a
H 0	20.00 ± 0.15	...
H $_2$	$19.21^{+0.13}_{-0.12}$	$\langle f_{\text{H}_2} \rangle = 0.24$
CO	14.20 ± 0.09	$\log \text{CO}/\text{H} = -5.92$
Zn $^+$	13.02 ± 0.02	$[\text{Zn}/\text{H}] = +0.34 \pm 0.12^b$
S $^+$	15.39 ± 0.06	$[\text{S}/\text{H}] = +0.15 \pm 0.13^b$
Fe $^+$	14.57 ± 0.01	$[\text{Fe}/\text{H}] = -1.05 \pm 0.12^b$
Si $^+$	15.15 ± 0.02	$[\text{Si}/\text{H}] = -0.48 \pm 0.12^b$
Ni $^+$	13.48 ± 0.03	$[\text{Ni}/\text{H}] = -0.86 \pm 0.12^b$

Notes. ^(a) Abundances are given considering the total neutral hydrogen column density $N(\text{H}) = N(\text{H}^0) + 2N(\text{H}_2)$. ^(b) Relative to solar abundances (Asplund et al. 2009).

super-solar with $[\text{Zn}/\text{H}] = +0.34$ and $[\text{S}/\text{H}] = +0.15$. Other species are depleted ($[\text{Fe}/\text{Zn}] = -1.39$, $[\text{Si}/\text{Zn}] = -0.82$). This indicates that a significant fraction of the refractory species is locked in solid phase onto dust grains. The presence of dust and its consequences are discussed in more detail in Sect. 4.2.

3.2.2. Neutral carbon

Because the ionisation potential of neutral carbon (C^0) is similar to the energy of the photons that destroy H_2 , C I is usually a good tracer of the presence of H_2 (Srianand et al. 2005). The expected positions of several C I lines usually fall out of the Lyman- α forest. We therefore initiated a programme to search for molecules along QSOs selected upon the presence of C I, as seen in the low-resolution SDSS spectra. Because of this selection, it is not surprising to detect strong C I lines in the UVES spectrum of J 1237+0647. The profile of C I absorption lines is complex and results from the blending of absorption lines from different components seen in different excitation levels (ground state: $^3\text{P}_0$, first excited level: $^3\text{P}_1$ and second excited level: $^3\text{P}_2$). Nevertheless, the high signal-to-noise and high spectral resolution allow us to clearly identify eight components. Most of them are also detected in the first excited level, while only the strongest two are detected in the second excited level. The fit to C I lines is shown in Fig. 4, with the corresponding parameters given in

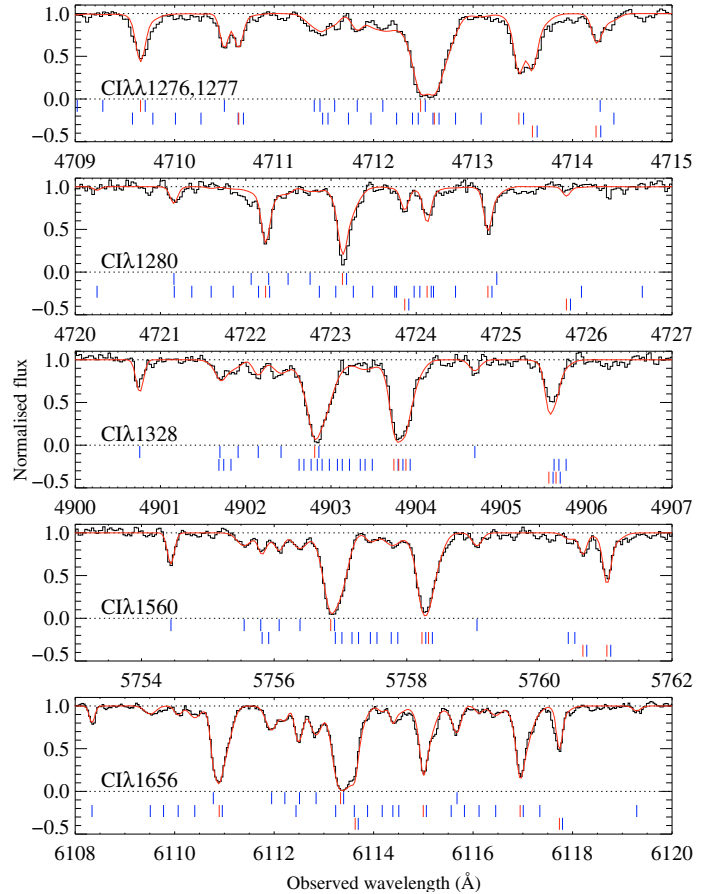


Fig. 4. Portions of the J 1237+0647 UVES spectrum around C I absorption lines. The short vertical marks in each panel represent the positions of absorption lines from the ground state ($^3\text{P}_0$), the first excited level ($^3\text{P}_1$) and the second excited level ($^3\text{P}_2$) from top to bottom, respectively. The component where CO is detected (at $z_{\text{abs}} = 2.68957$) is marked in red.

Table 4. We considered all optically thin absorption lines but did not include in the fit weak absorption lines in the region around C I $\lambda 1277$ where the placement of the continuum was uncertain. The main uncertainty in determining the C^0 column densities

Table 4. Column densities of neutral carbon in fine-structure levels.

z_{abs}	Δv^a (km s ⁻¹)	b (km s ⁻¹)	log $N(\text{C}^0)$ (cm ⁻²)			Total
			³ P ₀	³ P ₁	³ P ₂	
2.68802	-126	1.2 ± 0.1	13.24 ± 0.04	12.80 ± 0.03		13.37 ± 0.03
2.68872	-69	11.8 ± 0.6	12.84 ± 0.03	12.70 ± 0.05		13.08 ± 0.02
2.68888	-56	3.3 ± 1.3	12.37 ± 0.08			12.37 ± 0.08
2.68906	-41	3.0 ± 0.4	12.82 ± 0.02	12.56 ± 0.04		13.01 ± 0.02
2.68926	-25	7.4 ± 0.4	12.94 ± 0.02	12.87 ± 0.02		13.21 ± 0.01
2.68956	-1	1.4 ± 0.1	14.67 ± 0.04	14.46 ± 0.03	13.64 ± 0.02	14.90 ± 0.03
2.68960	+2	8.8 ± 0.1	13.98 ± 0.01	13.79 ± 0.01	12.98 ± 0.03	14.22 ± 0.01
2.69097	+114	3.6 ± 0.5	12.70 ± 0.02	12.42 ± 0.08		12.88 ± 0.03

Notes. ^(a) Relative to $z_{\text{abs}} = 2.68957$.

comes from the uncertainties on the oscillator strengths. As in previous works from our group (e.g. [Noterdaeme et al. 2007a,b](#); [Srianand et al. 2008b](#)) and others in the field ([Jorgenson et al. 2009](#)), we used the f -values from [Morton \(2003\)](#). Using the oscillator strengths from [Jenkins & Tripp \(2001\)](#) results in column densities that are 2 to 3 times lower.

The relative populations of the fine-structure levels of neutral carbon depend on the gas pressure. Because the kinetic temperature of the gas can be derived from the relative populations of the low rotational levels of H₂ (see Sect. 3.3), it is possible to measure the volumic density of the gas. From Fig. 2 of [Silva & Viegas 2002](#) (see also [Srianand et al. 2000](#)) and taking into account excitation by collisions and by the Cosmic Microwave Background radiation, we can see that the measured ratios $\log N(\text{C}^0, J = 1)/N(\text{C}^0, J = 0) = -0.2$ and $\log N(\text{C}^0, J = 2)/N(\text{C}^0, J = 0) = -1$ in the main component, which coincides with the position of the CO component, correspond to a volumic density of the order of $n_{\text{H}^0} \sim 50\text{--}60 \text{ cm}^{-3}$ for $T \sim 110 \text{ K}$ (see Sect. 3.3). Other components have similar fine-structure ratios, which indicate a similar thermal pressure. The kinetic temperature is probably higher in all other components (which is verified at least for the two other H₂-bearing components, see Table 6), which implies lower densities (e.g. $n_{\text{H}^0} \sim 1\text{--}10 \text{ cm}^{-3}$ in the component at $v = -127 \text{ km s}^{-1}$).

3.2.3. Neutral sulphur

Because the first ionisation potential of sulphur is 10.36 eV, this element is usually observed in its first ionised state (S⁺) in damped Lyman- α systems. In turn, sulphur is expected to be found in neutral form (S⁰) inside molecular clouds, where the surrounding UV field has been strongly attenuated. To date, S⁰ has been detected only in QSO absorbers where CO absorption is seen as well: in the systems at $z = 2.42$ towards SDSS J143912.04+111740.5 (hereafter J 1439+1117; [Srianand et al. 2008b](#)) and at $z = 1.64$ towards SDSS J160457.50+220300.5 ([Noterdaeme et al. 2009a](#)).

Absorption lines from eight transitions of neutral sulphur are detected in the UVES spectrum of J 1237+0647. This strengthens the claim that the presence of neutral sulphur flags molecular gas. Just like the presence of neutral carbon is a good indicator of that of H₂ ([Srianand et al. 2005](#), albeit with generally modest molecular fractions, [Noterdaeme et al. 2008a](#)), S I lines might well indicate the presence of CO. This is however of little practical use for pre-selecting CO-bearing DLAs from the low resolution SDSS spectra, because S I and CO lines are located in the same spectral region and have similar strengths.

We used all detected S I lines to constrain the column density and b parameter. Two components are needed to properly fit

Table 5. S⁰ column densities.

z_{abs}	Δv^a (km s ⁻¹)	log $N(\text{S}^0)$ (cm ⁻²)	b (km s ⁻¹)
2.68953	-3	12.29±0.06	1.1±0.5
2.68958	+1	13.19±0.04	0.7±0.1

Notes. ^(a) Relative to $z_{\text{abs}} = 2.68957$.

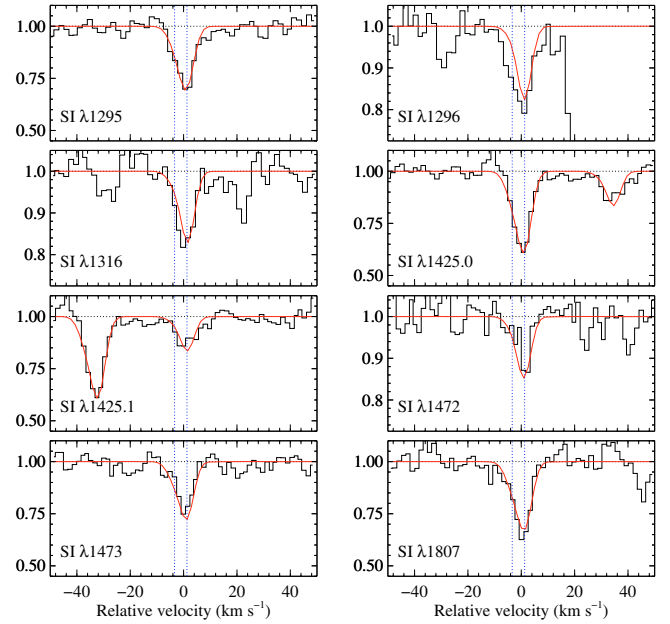


Fig. 5. Fit to S I absorption lines. Results of the best model fit with two components is overplotted ($\chi^2_{\nu} = 0.97$).

the data, with resulting $\chi^2_{\nu} \simeq 1$. The b -value obtained is less than 1 km s^{-1} for the main component, i.e., well below the spectral resolution ($\sim 6 \text{ km s}^{-1}$). However, b is well constrained, thanks to the relatively wide range that is spanned by the oscillator strength values. However, the fit is sensitive to the exact value of the spectral resolution. Therefore, to add confidence to the b and N measurements, we built the curve of growth for the detected S I lines, which does not depend on the spectral resolution (see Fig. 6). The error on the equivalent width measurements are conservative and take into account uncertainties in the continuum placement. From this figure, we confirm the small Doppler parameter. The measured column density nicely matches the sum of the individual column densities in the two components derived from the Voigt-profile fitting.

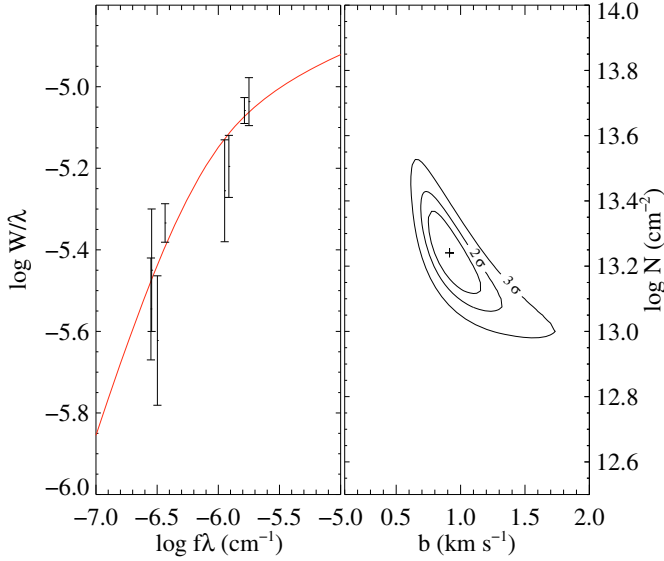


Fig. 6. Curve of growth analysis of S I absorption lines. *Left:* curve of growth. *Right:* confidence interval. Minimum χ^2_{ν} (0.91) is reached for $\log N(\text{S}^0) = 13.24$, $b = 0.91 \text{ km s}^{-1}$.

Table 6. H_2 column densities and excitation temperatures.

Component $z_{\text{abs}}, v \text{ (km s}^{-1}\text{)}$	$\log N(\text{H}_2, J)^a$ (cm^{-2})	b (km s^{-1})	T_{0-J} (K)
#1, $z = 2.68801, \Delta v = -127$	$16.28^{+0.10}_{-0.10}$		
$J = 0$	15.51 ± 0.05	3.3 ± 0.2	–
$J = 1$	16.08 ± 0.10	''	193^{+123}_{-54}
$J = 2$	15.39 ± 0.20	''	271^{+120}_{-63}
$J = 3$	15.13 ± 0.05	''	261^{+16}_{-15}
#2, $z = 2.68868, \Delta v = -73$	$17.62^{+0.08}_{-0.11}$		
$J = 0$	$16.94^{+0.20}_{-0.44}$	$0.4^{+0.8}_{-0.1}$	–
$J = 1$	17.51 ± 0.05	1.2 ± 0.1	≥ 117
$J = 2$	15.25 ± 0.03	3.3 ± 0.2	93^{+20}_{-9}
$J = 3$	14.89 ± 0.02	4.9 ± 0.2	131^{+19}_{-8}
#3, $z = 2.68955, \Delta v = -2$	$19.20^{+0.13}_{-0.12}$		
$J = 0$	18.65 ± 0.20	6.0 ± 0.1	–
$J = 1$	18.92 ± 0.10	''	108^{+84}_{-33}
$J = 2$	18.18 ± 0.10	''	190^{+66}_{-39}
$J = 3$	18.21 ± 0.10	''	252^{+52}_{-36}
$J = 4$	15.43 ± 0.10	7.9 ± 0.1	177^{+14}_{-11}
$J = 5$	14.95 ± 0.05	''	213^{+11}_{-10}

Notes. ^(a) The first line for each component gives the total H_2 column density in that component.

3.2.4. Neutral chlorine

Chlorine, with an ionisation potential of 12.97 eV, is a unique species among those that can be photoionised by photons with energy $h\nu < 13.6 \text{ eV}$. The dominant form of chlorine is Cl^+ when hydrogen is mostly in the atomic form. However, neutral chlorine (Cl^0) results from rapid exothermic ion-molecule reaction

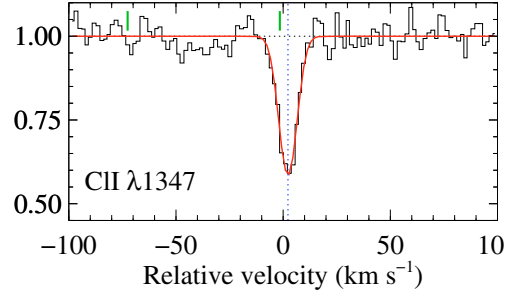


Fig. 7. CI I 1347 absorption at $z_{\text{abs}} = 2.68960$. A single component fit is overlotted.

between singly-ionised chlorine (Cl^+) and H_2 when H_2 is optically thick (Jura 1974). Therefore, the presence of neutral chlorine in the ISM is expected to be a good indicator of the presence of molecular gas.

Cl^0 is clearly detected at $z_{\text{abs}} = 2.689560$, associated with the strongest H_2 component. We measure $\log N(\text{Cl}^0) = 13.01 \pm 0.02$ from the fit to the CI I 1347 absorption line, with $b = 4.5 \pm 0.4 \text{ km s}^{-1}$ (see Fig. 7). From the non-detection of CI II 1071 we derive $\log N(\text{Cl}^+) < 13.4$ at the 3σ confidence level, which translates to $f_{\text{Cl}^0} \equiv N(\text{Cl}^0)/(N(\text{Cl}^0) + N(\text{Cl}^+)) > 0.3$. As the fraction of chlorine in neutral form is expected to follow approximately that of hydrogen in molecular form (Jura & York 1978; see also Sonnentrucker et al. 2002), the lower limit on f_{Cl^0} indicates that hydrogen could be mostly molecular at the place where we detect Cl^0 . We indeed show in the next section that the molecular fraction is particularly high in this component.

3.3. Molecular hydrogen

Molecular hydrogen is detected in three components at $z_{\text{abs}} = 2.68801, 2.68868$ and 2.68955 , spread over $\sim 125 \text{ km s}^{-1}$. The strongest component at $z_{\text{abs}} = 2.68955$ also features HD and CO absorption lines. The UVES spectrum covers numerous Lyman bands ($\text{B}^1\Sigma_u^+ \leftarrow \text{X}^1\Sigma_g^+$) as well as some Werner bands ($\text{C}^1\Pi_u^+ \leftarrow \text{X}^1\Sigma_g^+$), which allows for an accurate measurement of the H_2 column densities in each component and in different rotational levels. A portion of the UVES spectrum covering the H_2 Lyman (1–0) band is shown in Fig. 8, while the full velocity plots for different rotational levels are shown in Figs. 16 to 21. The measured column densities and corresponding excitation temperatures are given in Table 6. Below we refer to these components as #1, #2 and #3 from the bluest to the reddest.

Component #1 has a similar width in all rotational levels and requires a single Doppler parameter to describe the profiles of the H_2 $J = 0$ to $J = 3$ absorption lines. The Doppler parameter ($b \sim 3 \text{ km s}^{-1}$) is likely to be dominated by turbulent motions.

Component #2 presents a broadening of the profiles with increasing rotational level and requires different Doppler parameters. This behaviour has already been observed in the Galactic ISM (e.g. Jenkins & Peimbert 1997; Lacour et al. 2005b) but also in high-redshift damped Lyman- α systems (Noterdaeme et al. 2007a). Doppler parameters can be accurately measured even when they are significantly smaller than the spectral resolution thanks to the presence of numerous transitions with different oscillator strengths. However, the measurement of b in the first rotational level ($J = 0$) remains difficult because of the small number of unblended absorption lines (see Table 6). However, there are enough transitions from the $J = 0$ and 1 levels together

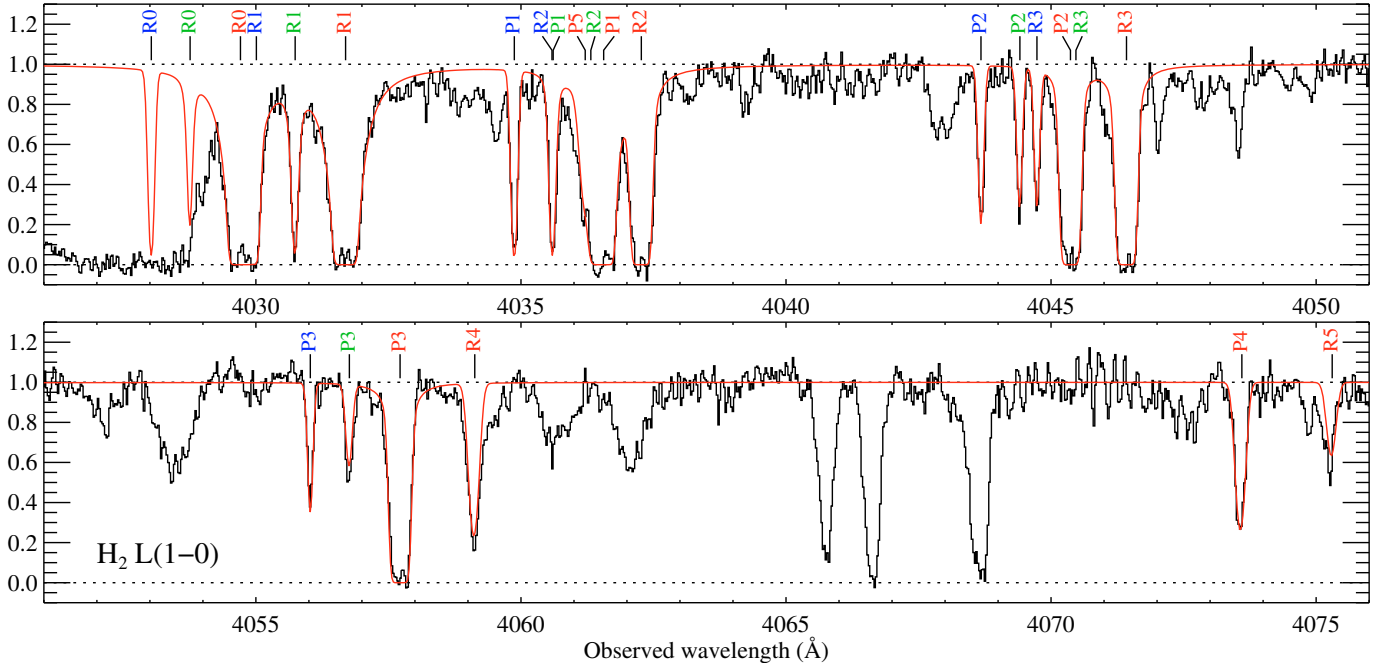


Fig. 8. Portion of the UVES spectrum of J1237+0647 covering the (1–0) Lyman band of H₂. The labels indicate the branches (“R”, “P” for $\Delta J = -1, +1$, respectively) and the rotational levels of the lower states. Absorptions from different components are indicated by different label colours (#1: blue, #2: green, #3: red).

to ascertain that the Doppler parameter of the transitions from these two levels is very small, of the order of $b \sim 1 \text{ km s}^{-1}$. This is consistent with thermal excitation with a temperature of $T_k \sim 120 \text{ K}$, which is similar to what is measured from T_{01} .

The reddest component (#3) is very strong, with damping wings seen for rotational levels up to $J = 3$, which allows for accurate measurement of the column densities. Non-saturated lines for $J = 4$ and 5 also allow for accurate measurements in these rotational levels. Component #3 is particularly interesting not only because of its high H₂ column density, but because it also contains deuterated molecular hydrogen, carbon monoxide as well as neutral sulphur and neutral chlorine, all of which have been very rarely detected at high redshift. Because the column density of H₂ is high, the $J = 0$ and $J = 1$ levels are self-shielded and the collisional time-scale is much shorter than the photo-dissociation time-scale. Therefore, the $N(J = 1)/N(J = 0)$ ratio is maintained at the Boltzmann equilibrium value. This means that T_{01} provides a robust measurement of the kinetic temperature. We measure $T_{01} = T_{\text{kin}} \sim 110 \text{ K}$, which is similar to the temperature in the local interstellar medium ($T_{\text{kin}} \sim 80 \text{ K}$, Savage et al. 1977).

We measure a total column density $\log N(\text{H}_2) = 19.21^{+0.13}_{-0.12}$ in the sub-DLA system with far the most important contribution coming from the CO-bearing component (component #3). This corresponds to a large mean molecular fraction $\langle f_{\text{H}_2} \rangle = 2N(\text{H}_2)/(2N(\text{H}_2) + N(\text{H}^0)) = 0.24^{+0.13}_{-0.08}$. However, the centre of the HI Lyman- α absorption line is clearly shifted from component #3 by more than 50 km s^{-1} (see Fig. 2). This means that the amount of atomic hydrogen in the CO-bearing cloud is much smaller than $\log N(\text{H}^0) = 20$, and the value given above should be considered as a lower limit on the actual molecular fraction in the CO-bearing cloud (i.e. $f_{\text{H}_2} > 1/4$). As noticed in the previous section from the presence of Cl⁰ absorption associated with #3, the molecular fraction in the component #3 is probably not far from unity.

Table 7. HD column densities.

Component	$\log N(\text{HD}, J) (\text{cm}^{-2})$	$b (\text{km s}^{-1})$
$z_{\text{abs}} = 2.68956, \Delta v = -1 \text{ km s}^{-1}$		
$J = 0$	14.48 ± 0.05	4.5 ± 0.2
$J = 1$	≤ 13.60	''

3.4. HD and the D/H ratio

Several Lyman-band HD lines from the first two rotational levels are detected in the UVES spectrum (see Fig. 9). Unfortunately, $J = 1$ lines are either severely blended or redshifted in regions of bad signal-to-noise ratio and we can only derive an upper limit on the column density for this rotational level. We measured the HD column density in the $J = 0$ rotational level from fitting the HD L5R0 and HD L8R0 lines, which are unblended. Measurements are summarised in Table 7.

We measured $N(\text{HD})/2N(\text{H}_2) = 0.95 \times 10^{-5}$. This is about 10 times higher than what is measured in the Galactic ISM (Lacour et al. 2005a) for $f = 0.24$. Because this ratio is known to increase with the molecular fraction (Lacour et al. 2005a), HD/2H₂ might provide a lower limit on D/H for $f_{\text{H}_2} < 1$. If, as discussed previously, the actual molecular fraction in the HD-bearing cloud towards J1237+0647 is higher than 0.25, both HD and H₂ could be self-shielded and $\text{HD}/2\text{H}_2 \sim \text{D}^0/\text{H}^0$. The value we obtain is then consistent with the D^0/H^0 ratio measured in the Galactic disc (Linsky et al. 2006). This corresponds to an astration factor of ~ 3 when comparing to the primordial value as measured in low-metallicity absorption systems ($\text{D}/\text{H} = 2.82 \pm 0.2 \times 10^{-5}$, Pettini et al. 2008; Ivanchik et al. 2010; see however Srianand et al. 2010) or derived from the baryon density parameter (Steigman 2007).

Note that all five high-redshift HD detections to date yield relatively high D/H values despite significant metal enrichment: $N(\text{HD})/2N(\text{H}_2) = 1.5 \times 10^{-5}, 3.6 \times 10^{-5}, 7.9 \times 10^{-5}$ and

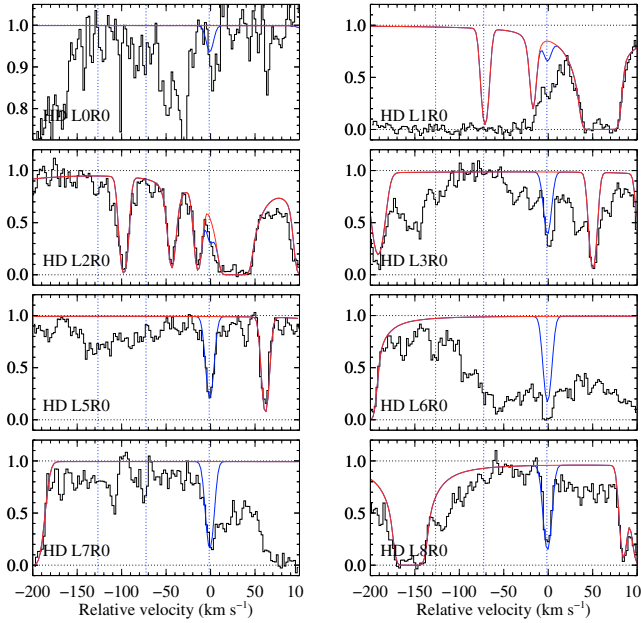


Fig. 9. Fit to HD $J = 0$ lines. The blue profile is the contribution from HD while the red profile represents H_2 . The velocity of the H_2 detected components are indicated by vertical dotted lines.

1.6×10^{-5} towards, J 1439+1117 (Noterdaeme et al. 2008b), Q 1232+082 (Ivanchik et al. 2010), J 2123-0500 and FJ 0812+32 (Tumlinson et al. 2010). Because deuterium is easily destroyed as interstellar gas is cycled through stars, large deuterium abundances are difficult to reproduce with closed-box models. However, these are well explained by models including infall of primordial gas (e.g. Prodanović & Fields 2008). If the velocity-metallicity correlation found by Ledoux et al. (2006a) is the consequence of an underlying mass-metallicity relation, we can expect that a high astration of deuterium in high metallicity systems is roughly compensated by a strong infall of primordial material onto massive galaxies. However, Tumlinson et al. (2010) noted that HD/ $2H_2$ ratios in high- z absorption systems lie in a narrow range well above the value measured in the Galaxy, while these systems present a broad diversity in terms of metallicities and molecular fractions. This puzzling behaviour led them to conclude that it could be premature to use the HD/ $2H_2$ ratio to derive Ω_b , given our actual understanding of interstellar chemistry. In addition, we note that for QSO absorbers, we only have access to the properties of the gas (metallicities, molecular fractions) averaged over the line of sight. These may not be representative of the actual chemical abundances in the HD-bearing cloud. Indeed, only the total $N(H^0)$ can usually be measured and the metal components are blended into a smooth absorption profile. It is therefore necessary to be careful and to study each system in detail (e.g. Balashev et al. 2010) to derive the local chemical and physical conditions in the cloud.

3.5. Carbon monoxide

Absorptions from eight $A^1\Pi(v') \leftarrow X^1\Sigma^+(v = 0)$ bands of CO (from AX (0–0) to AX (7–0)), the $C^1\Sigma(v' = 0) \leftarrow X^1\Sigma^+(v = 0)$ Rydberg band and the $d^3\Delta(v' = 5) \leftarrow X^1\Sigma^+(v = 0)$ inter-band system are detected at $z_{\text{abs}} = 2.68957$ towards J 1237+0647, see Fig. 10. Each band produces a complex absorption profile, which results from superposition of absorption lines from different rotational levels in the P and R branches. The resolving

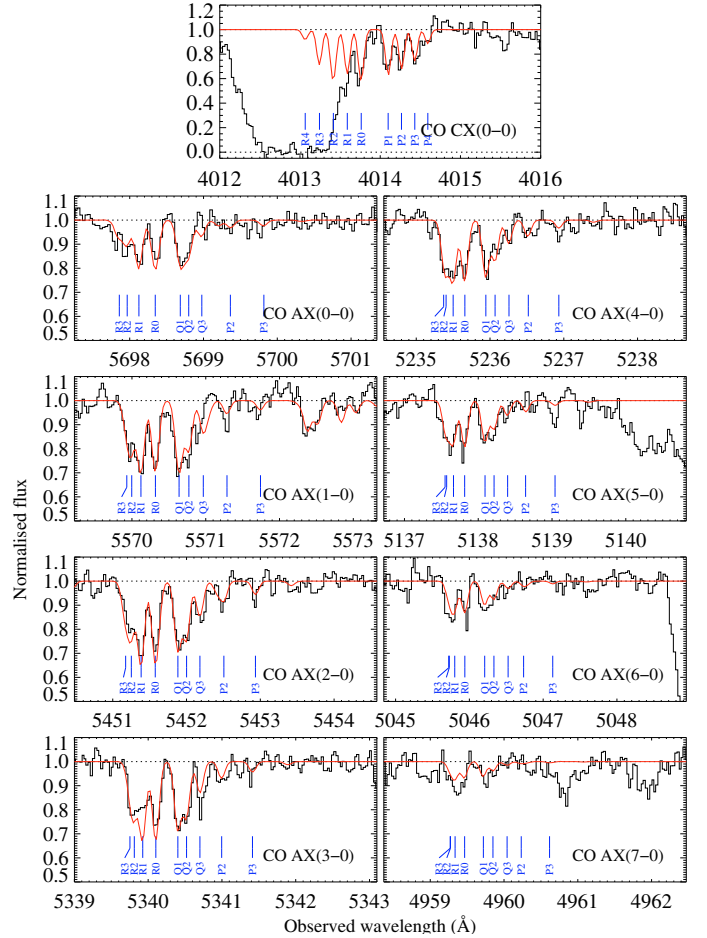


Fig. 10. Fit to CO lines ($\chi^2_v = 1.1$). The additional absorption seen in the CO AX(1–0) panel is owing to the CO dX(5–0) inter-band system. Short vertical lines indicate the branch (“R”, “P” or “Q”) and the rotational level of the transition.

power ($R \sim 50\,000$) and the signal-to-noise ratio ($SNR \sim 28$) of the UVES spectrum are high enough to individually measure the column densities in rotational levels up to $J = 3$. In addition, the $J = 4$ level is probably detected in the P branch of CX (0–0). We used the AX and dX bands that fall outside of the Lyman- α forest to measure the column densities in rotational levels up to $J = 3$. The AX (3–0) band is affected by a spike that is likely due to a cosmic ray impact at the position of the R branch, and this region is therefore not considered when fitting the profile. In addition, we use the CX(0–0) band up to $J = 4$. This band is the strongest one available but the corresponding rest wavelength (1088 Å) makes it redshifted in the Ly- α forest. Fortunately, only the R branch is blended but the P branch is free from any blend. Moreover, the CX rotational levels are well separated at the UVES spectral resolution.

The results of the fit are presented in Table 8. Two errors are quoted in this table for the column densities: the first one is the rms error on the fit, while the second one reflects the uncertainties resulting from the continuum placement. The latter uncertainties were estimated by changing the normalisation by plus or minus 0.5σ (i.e. about $\pm 2\%$ for $SNR = 28$) around the best continuum fit. The total CO column density we derive is $\log N(\text{CO}) = 14.17 \pm 0.09$, implying a high CO to H_2 ratio of $N(\text{CO})/N(\text{H}_2) = 10^{-5}$. This is typical of what is seen in translucent clouds (see Burgh et al. 2010).

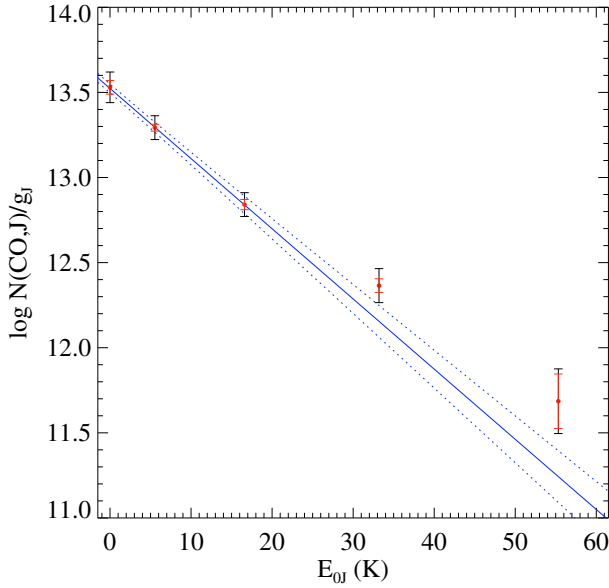


Fig. 11. Excitation diagram of CO rotational levels. Errors on the column densities from fitting the lines are represented by the small red error bars while the long black error bars take into account the uncertainty in the continuum placement. The plain line represents the linear regression fit using $J = 0$ to $J = 2$ measurements and the best continuum fit. 1σ errors are represented by dashed lines.

Table 8. CO column densities and excitation temperatures.

Component	$\log N(\text{CO}, J)^a$	b (km s $^{-1}$)	T_{0-J}^b (K)
$z_{\text{abs}} = 2.68957$	14.17 ± 0.09^c	0.9 ± 0.1	
$J = 0$	$13.53 \pm 0.04 \pm 0.08$	"	–
$J = 1$	$13.77 \pm 0.02 \pm 0.07$	"	$10.1^{+4.3}_{-2.0}$
$J = 2$	$13.54 \pm 0.03 \pm 0.06$	"	$10.5^{+1.4}_{-1.1}$
$J = 3$	$13.21 \pm 0.04 \pm 0.09$	"	$12.4^{+1.0}_{-0.9}$
$J = 4$	$12.64 \pm 0.16 \pm 0.11$	"	$13.0^{+1.9}_{-1.6}$

Notes. ^(a) Quoted errors on column densities are errors from fitting the Voigt profiles and errors owing to continuum placement respectively. The latter were estimated by varying the continuum by $\pm 0.5\sigma$. ^(b) The errors on T_{0-J} represent the extremum values for the different sets of continuum. ^(c) Total CO column density.

In Fig. 11 we show the excitation diagram of CO. It is clear that the population of the first three rotational levels can be reproduced with a single excitation temperature. We measure this excitation temperature by performing a linear fit of $\log N(\text{CO}, J)/g_J$ vs. the energy of the levels (E_{0J}). The fit and 1σ range in Fig. 11 corresponds to the best-fit continuum. In order to estimate the effect of the continuum placement, we repeated the linear fit for each set of continua and took the extrema as representative of the range of possible values for $T_{\text{ex}}(\text{CO})$. This gives $T_{\text{ex}} = 10.5^{+0.8}_{-0.6}$. Note that the effect of the continuum placement is mainly a change in the total CO column density, while little change on the slope of the linear fit (i.e. the excitation temperature). The CO excitation temperature is well below the kinetic temperature of the gas. This means that the gas density is low enough so that radiative processes are likely to dominate the downwards cascade, as predicted in diffuse molecular and translucent clouds (e.g. Warin et al. 1996). Indeed, the population ratios of the neutral carbon fine-structure levels indicate a

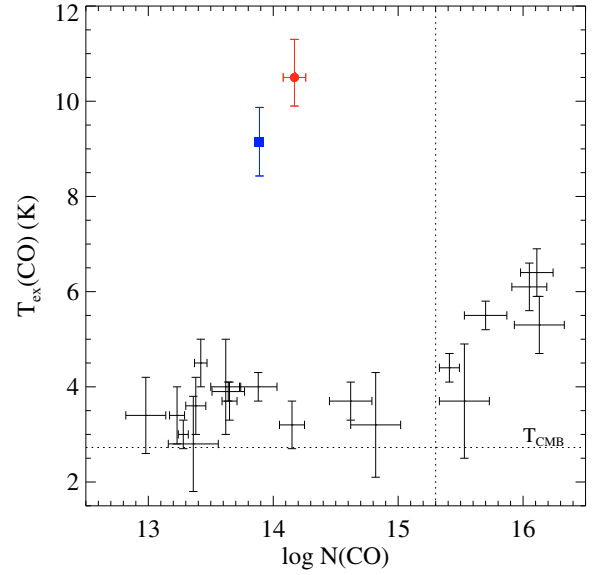


Fig. 12. Excitation temperature of CO as a function of the total CO column density. Black error bars are measurements at $z = 0$ from Burgh et al. (2007). The red filled circle with error bars is our measurement at $z = 2.7$ towards J 1237+0647 while the blue square represents the measurement at $z = 2.4$ towards J 1439+1117 (Srianand et al. 2008b). Only a large range on T_{ex} (6–16 K) could be determined for the system at $z = 1.64$ towards SDSS J160457.50+220300.5 (Noterdaeme et al. 2009a), and is therefore not represented in the figure. The vertical dotted line is indicative of a change of behaviour: below this limit there is no correlation between $N(\text{CO})$ and $T_{\text{ex}}(\text{CO})$.

volumic density of the order of $n_{\text{H}^0} \sim 50 \text{ cm}^{-3}$, which is well below the critical density at which the collisional de-excitation rate of $\text{CO}(J = 1-0)$ equals that of the spontaneous emission ($n_{\text{crit}} \sim 1000 \text{ cm}^{-3}$, Snow & McCall 2006). Indeed, in terms of density and molecular fraction, the CO-bearing system presented here is very similar to that presented in Srianand et al. (2008b) where we concluded that collisional excitation of CO is negligible. Note however that the fine-structure levels of C^0 only give the average volumic density. The actual local volumic density in the CO-bearing cloud could be higher. A small shift ($\sim 1 \text{ km s}^{-1}$) is measured between the strongest CI feature and the CO component. This may indicate that the two species are not completely co-spatial.

From the radiative code RADEX (van der Tak et al. 2007) we expect the excitation temperature of CO to be about one degree higher than the expected temperature of the cosmic microwave background (CMB) radiation ($T_{\text{CMB}}(z = 2.69) = 10.05 \text{ K}$) as soon as the collision partner (H^0 , H_2 and He) density is higher than 50 cm^{-3} . This explains that the excitation temperature we measure is slightly higher than what is expected from excitation by the CMB radiation alone.

In Fig. 12 we compare the excitation temperature of CO at high redshift with that in the local Universe. In the local ISM, the temperature is seen a few degrees above T_{CMB} at low CO column densities and rises for column densities above $N(\text{CO}) = 2 \times 10^{15} \text{ cm}^{-2}$. This is due to the increased importance of photon trapping at higher column densities (Burgh et al. 2007). The values observed at high redshift are significantly higher than the local ones, despite similar $N(\text{CO})$ and kinetic temperatures. This clearly means that the main physical difference between high redshift and local lines of sight is the higher CMB temperature at high redshift. This provides a strong positive test of the hot Big-Bang theory. Another consequence of Fig. 12 is that only

CO-bearing systems with $\log N(\text{CO}) < 15$ – for which there is no correlation between $N(\text{CO})$ and $T_{\text{ex}}(\text{CO})$ – are good places to measure the evolution of T_{CMB} with cosmic time.

Interestingly, although the differences are small and within errors, we measure a systematic trend, $T_{04} \gtrsim T_{03} \gtrsim T_{02} \gtrsim T_{01}$, regardless of the exact continuum placement. This indicates that while CMB photons dominate the rotational excitation of CO, other mechanisms are at play. We fail to reproduce the increasing temperature with increasing rotational level with RADEX. However, this behaviour has already been noticed in the local ISM (Sonnentrucker et al. 2007; Sheffer et al. 2008) and could be explained by the selective self-shielding of low rotational lines for $\log N(\text{CO}) > 14$ (Warin et al. 1996). The self-shielding of far-UV Rydberg bands of CO (those relevant to the photo-destruction process) could be more effective than previously thought (Sheffer et al. 2003). In addition, the presence of H₂ lines in the same spectral region can contribute to an effective shielding of CO lines. Finally, radiative pumping from CO emission lines due to nearby dense molecular clouds could contribute to populate the higher rotational levels in the absorbing cloud (Wannier et al. 1997). If the increasing temperature with increasing rotational level is physical, then $T_{01} = 10.1^{+4.3}_{-2.0}$ K could better represent the excitation by the CMB alone.

4. The nature of the absorbing cloud

4.1. Summary of the physical properties in the CO component

In the previous sections we derived physical properties of the gas associated to the molecular absorptions seen at $z_{\text{abs}} \sim 2.68957$ towards J 1237+0647.

From the analysis of C⁰, we derived that the molecular fraction, $2N(\text{H}_2)/(2N(\text{H}_2)+N(\text{H}^0))$, in the CO component is larger than 1/4 for a super-solar metallicity: $Z(\text{Zn}, \text{S}) = +0.34, +0.15$. From the populations of the C⁰ ground-state fine structure levels, we found that the particle density is of the order of $\sim 50 \text{ cm}^{-3}$. The analysis of the H₂ rotational levels yields a kinetic temperature of ~ 100 K and CO is mainly excited by radiative interaction with the CMBR.

We can have an indication of the electronic density in the cloud thanks to the S⁰/S⁺ ratio. Assuming that the mean ratio can be derived using the column densities in the strongest components we measure: $\log N(\text{S}^0)/N(\text{S}^+) = -1.72$. The electronic density, n_e , is derived from the ionisation equilibrium between the two species

$$\Gamma n(\text{S}^0) = \alpha n_e n(\text{S}^+), \quad (1)$$

where Γ is the photoionisation rate of S⁰ and α the combination rate of S⁺. Taking the ratio in diffuse gas of the Galaxy ($\Gamma/\alpha \sim 95 \text{ cm}^{-3}$, Péquignot & Aldrovandi 1986; see also Table 8 of Welty et al. 1999b) we derive $n_e \sim 1.85 (\Gamma/\alpha/95) \text{ cm}^{-3}$. Note that this is an average value in the strongest metal component. Because S⁰ and S⁺ are not co-spatial – as indicated by the different Doppler-parameters – the electron density derived here should be considered as an upper limit in the molecular gas.

We could apply the same reasoning to carbon, but we cannot measure the C⁺ column density because C II $\lambda 1334$ is highly saturated. In turn we can derive $N(\text{C}^+)$ from the sulphur electronic density measurement. We find

$$\frac{N(\text{C}^+)}{N(\text{S}^+)} = \frac{\Gamma(\text{C}^0)/\alpha(\text{C}^+) N(\text{C}^0)}{\Gamma(\text{S}^0)/\alpha(\text{S}^+) N(\text{S}^0)}. \quad (2)$$

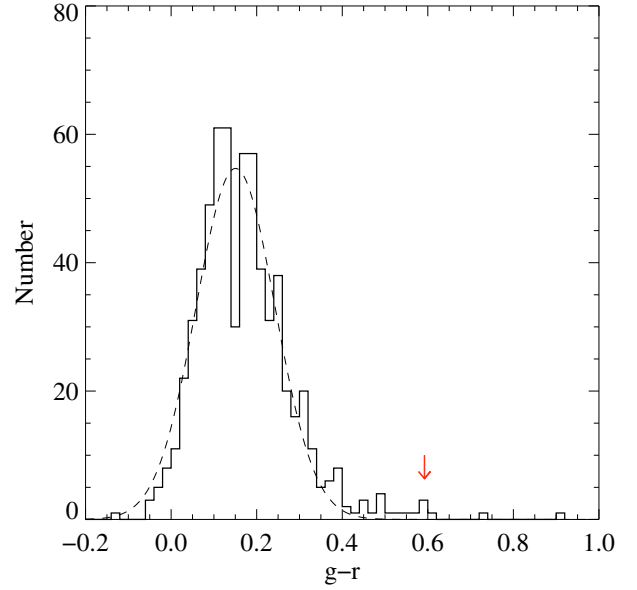


Fig. 13. Distribution of SDSS photometric $g - r$ values for a sample of 650 non-BAL QSOs at $2.681 < z < 2.881$. The arrow indicates the position of J 1237+0647.

Using $\Gamma(\text{C}^0)/\alpha(\text{C}^+) = 32 \text{ cm}^{-3}$ yields $\log N(\text{C}^+)/N(\text{S}^+) = 1.26$, which is close to the solar ratio.

4.2. Dust content

Figure 13 shows the distribution of $g - r$ colours for 650 non-BAL quasars with redshifts similar to J 1237+0647 ($\Delta z = 0.1$). The median $g - r$ value of the distribution is 0.16 with a standard deviation of 0.11. This dispersion reflects the colour variation from one QSO to the other and is not due to uncertainties in the SDSS photometric measurements (which are better than 0.03 mag). This implies that the colour excess of J 1237+0647 ($g - r = 0.59$) is significant at the 3.9σ level. This significance increases to 4.9σ if we consider a Gaussian fit to the distribution. Indeed, there is a tail towards large colour excesses, which shows the existence of a population of reddened quasars. J 1237+0647 is among the reddest 1.5% quasars with $g - r > 0.5$.

The observed flux-calibrated SDSS spectrum of J 1237+0647 is matched by the QSO composite spectrum from Vanden Berk et al. (2001) reddened by a SMC-like extinction law with $E(B - V) = 0.05 \pm 0.01$. The spectrum together with the fit is shown in Fig. 14. Note that because we are only interested in the slope of the continuum of J 1237+0647 and the strength of emission lines varies strongly from one quasar to the other, we do not include emission line regions in our fit. Absorption lines were also ignored during the fit as well as the wavelength range bluewards of the Ly- α emission.

In order to further estimate the probability that the above reddening might be due to a peculiar intrinsic QSO shape, we used a technique described in Srianand et al. (2008a) and Noterdaeme et al. (2009a): we repeated the spectral slope fitting assuming an absorber at $z = 2.69$ for a control sample of 82 non-BAL QSOs with similar emission redshifts ($2.76 < z_{\text{em}} < 2.80$) and spectra with i -band S/N ratios higher than 5. We find that the continuum slope of J 1237+0647 deviates at the 98% confidence level from the mean slope of other quasars. Below we will therefore consider that the colour excess of J 1237+0647 is owing to the presence of dust at $z_{\text{abs}} = 2.69$.

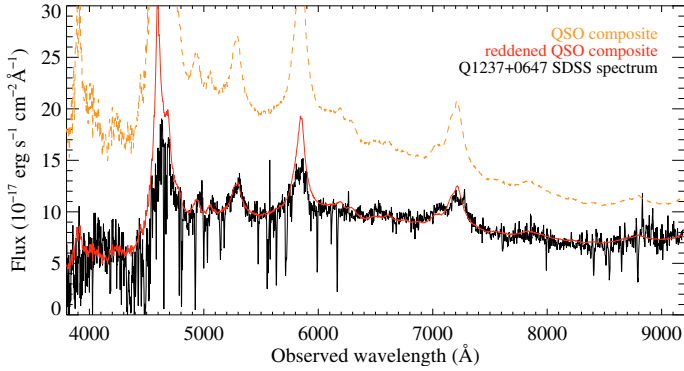


Fig. 14. SDSS spectrum of J1237+0647 (black). The orange dashed spectrum is the unreddened SDSS QSO composite spectrum (Vanden Berk et al. 2001). The red spectrum is the composite spectrum reddened by a SMC extinction law with $E(B - V) = 0.053$ (i.e. $A_V = 0.14$, in the rest-frame of the absorber). The fit was performed using an IDL code based on MPFIT (Markwardt 2009).

Our best-fit model predicts $J = 16.9$, $H = 16.3$, $K = 15.5$ (with typical errors of 0.2 mag), while the observed Two Micron All Sky Survey (2MASS) magnitudes are $J = 16.8 \pm 0.2$, $H = 15.7 \pm 0.2$, and $K > 15.0$. The agreement is reasonable although not perfect in the H-band. Indeed, the 2MASS magnitudes come from *Point Source Reject Table* for objects with very low SNR. These measurements are known to suffer from flux overestimation, which can easily explain the discrepancy between predicted and measured H-band magnitudes. In addition, the presence of the H- β emission line in the H-band increases the uncertainty of our estimate. This together with the fact that the 2MASS and SDSS observations were taken five years apart, makes our predicted magnitudes consistent with the 2MASS data. Unfortunately, there are no SDSS measurements at different epochs for this object to monitor any variation in the QSO flux.

The measured reddening, although significant, is marginally higher than what is seen in the general population of DLAs ($E(B - V) \sim 0.04$, Ellison et al. 2005). Interestingly, the integrated extinction-to-gas ratio measured towards J1237+0647, $A(V)/N(\text{H}^0) = 1.5 \times 10^{-21} \text{ mag cm}^2$ is 20 times higher than the average value for the SMC ($7.5 \times 10^{-23} \text{ mag cm}^2$, Gordon et al. 2003) and about 50 times higher than the mean value measured in high redshift DLAs ($2\text{--}4 \times 10^{-23} \text{ mag cm}^2$, Vladilo et al. 2008). This means that if the H^0 column density were higher, the extinction induced would have been so high that the QSO would have been missed by the SDSS target selection². Note that the moderate extinction $A_V = 0.14$ in the rest-frame of the absorber, already produces an extinction of nearly 1 mag in the g-band. This further supports the possibility that current surveys can miss a large number of cold clouds (Noterdaeme et al. 2009a). If, as discussed before, a large fraction of the atomic hydrogen is not associated with the molecular component, then the extinction-to-gas ratio in the molecular component is even higher. The line-of-sight to J1237+0647 probably passes through a relatively thin slab of dusty gas. This is also supported by the high depletion factors measured in the CO-bearing component ($[\text{Fe}/\text{Zn}] \sim -2.0$, $[\text{Ni}/\text{Zn}] \sim -1.6$, $[\text{Si}/\text{Zn}] \sim -1.7$, $[\text{Cr}/\text{Zn}] < -1.6$). This abundance pattern is typical of what is seen in cold gas of the Galactic disc (Welty et al. 1999a).

² For $\log N(\text{H}^0) = 20.65$ and same extinction-to-gas ratio, the predicted magnitude already reaches the $i = 20.2$ limit set by the SDSS-II collaboration for spectroscopy of high redshift quasars.

If we consider the extinction only, then the *sightline* studied in this paper is not translucent. Indeed, the historical definition of translucent corresponds to $1 < A_V < 5$. However, as noted by Rachford et al. (2002), even translucent sightlines can result from the concatenation of multiple diffuse clouds along the line-of-sight. This kind of scenario has already been advocated to explain the bimodal distribution in the $\log N(\text{H}_2)$ distribution for a given column density of dust (Noterdaeme et al. 2008a). Therefore extinction may not be the best parameter to define translucent clouds. Rachford et al. (2002) and Snow & McCall (2006) have proposed that the definition of translucent *clouds* should be based on the local properties of the gas rather than on the integrated properties along the line-of-sight. Following this suggestion, Burgh et al. (2010) have presented a definition based on the abundance of hydrogen and carbon in molecular forms. These authors noted that with this definition, translucent clouds do not necessarily produce strong reddening.

4.3. The molecular fractions of H and C

In Fig. 15 we plot the ratios $N(\text{CO})/N(\text{H}_2)$ and $N(\text{CO})/N(\text{C}^0)$ versus the hydrogen molecular fraction, f_{H_2} , for the two systems towards J1237+0647 (this work) and J1439+1117 (Srianand et al. 2008) together with measurements in the ISM of the Galaxy (Burgh et al. 2010). The latter authors define translucent clouds as clouds with $\text{CO}/\text{H}_2 > 10^{-6}$ and $\text{CO}/\text{C}^0 > 1.0$ for $f_{\text{H}_2} > 0.4$. In the cloud toward J1237+0647, we find $N(\text{CO})/N(\text{H}_2) = 10^{-5}$ for $\langle f_{\text{H}_2} \rangle = 0.24$. However, as discussed previously, the hydrogen molecular fraction in the CO-bearing cloud is higher than this. In addition, the amount of carbon in CO molecules is about that in atomic form³. We therefore conclude that the *cloud* in front of J1237+0647 is indeed a translucent cloud, an ideal laboratory for probing astrochemistry at high redshift.

5. Conclusion

From our VLT survey for H_2 in DLAs it appears that neutral carbon is generally observed in the same components that feature H_2 (Srianand et al. 2005; Noterdaeme et al. 2008a). We therefore selected the rare SDSS lines of sight in which C I absorptions are present. From UVES follow-up observations we have detected strong absorptions from H_2 , HD and CO along SDSS J123714.60+064759.5. This is a beautiful and peculiar case where detailed analysis of the physical properties of the gas is possible.

The H^0 column density is low, $\log N(\text{H}^0) = 20.00 \pm 0.15$ and corresponds to what is usually called a sub-DLA (Péroux et al. 2002). The corresponding overall metallicity is super solar with $[\text{Zn}/\text{H}] = +0.34$ and $[\text{S}/\text{H}] = +0.15$. The system features three H_2 components that span $\sim 125 \text{ km s}^{-1}$, the strongest of which, with $\log N(\text{H}_2) = 19.20$, does *not* coincide with the centre of the H I absorption. This means that the molecular fraction in this component is higher than the mean molecular fraction $\langle f_{\text{H}_2} \rangle = 1/4$ in the system.

From the populations of the low H_2 rotational levels we measured the kinetic temperature of the gas to be around 100 K in the strongest component, where HD and CO are also detected. The detection of S^0 and C^0 implies that the gas is shielded from the surrounding far-UV radiation field.

³ Note that we use here $\log N(\text{C}^0, J = 0, 1, 2) = 14.39 \pm 0.10$ derived using f -values from Jenkins & Tripp (2001) to enable comparison with the work by Burgh et al. (2010).

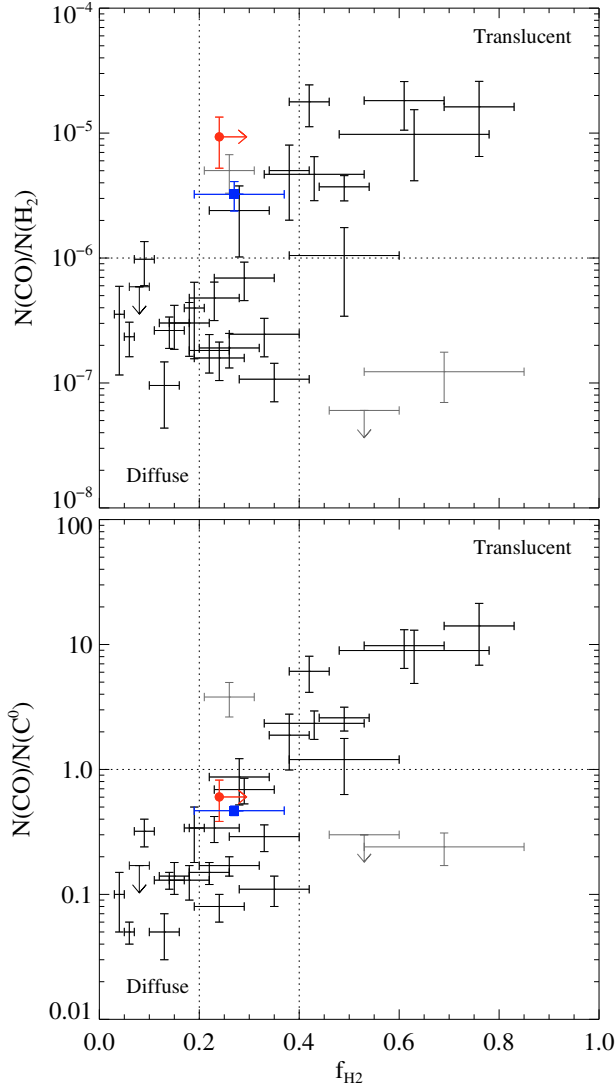


Fig. 15. $N(\text{CO})/N(\text{H}_2)$ and $N(\text{CO})/N(\text{C}^0)$ versus the molecular fraction, f_{H_2} . The red filled circle with error bars is our measurement at $z = 2.7$ towards J 1237+0647, while the blue square represents the measurement at $z = 2.4$ towards J 1439+1117 (Srianand et al. 2008b). Other points represent measurements in the Galactic ISM with three outliers represented in grey (Burgh et al. 2010). These are peculiar systems with uncertain measurements by the same authors. The horizontal dotted line marks the limit between diffuse and translucent clouds, while the two vertical dotted lines mark the transition range between these two regimes. Note that the molecular fraction we indicates toward J 1237+0647 is a lower limit. In order to facilitate the comparison with local values, we use $N(\text{C}^0)$ measured with the f -values from Jenkins & Tripp (2001).

The relative populations of the C^0 fine structure levels yield an estimate of the average hydrogen density in the main component of about $50\text{--}60\text{ cm}^{-3}$. At these densities collisions are not frequent enough to dominate the rotational excitation of CO molecules, and radiative processes are likely to determine the CO rotational populations. The excitation temperature we measure ($T_{\text{ex}}(\text{CO}) = 10.5\text{ K}$) is significantly higher than that measured in the local ISM ($T_{\text{ex}}(\text{CO}) \sim 3.5\text{ K}$) for similar CO column densities, molecular fractions, and kinetic temperatures. We show that the higher rotational excitation of CO towards J 1237+0647 results from the higher temperature of the cosmic microwave background at the redshift of the absorber,

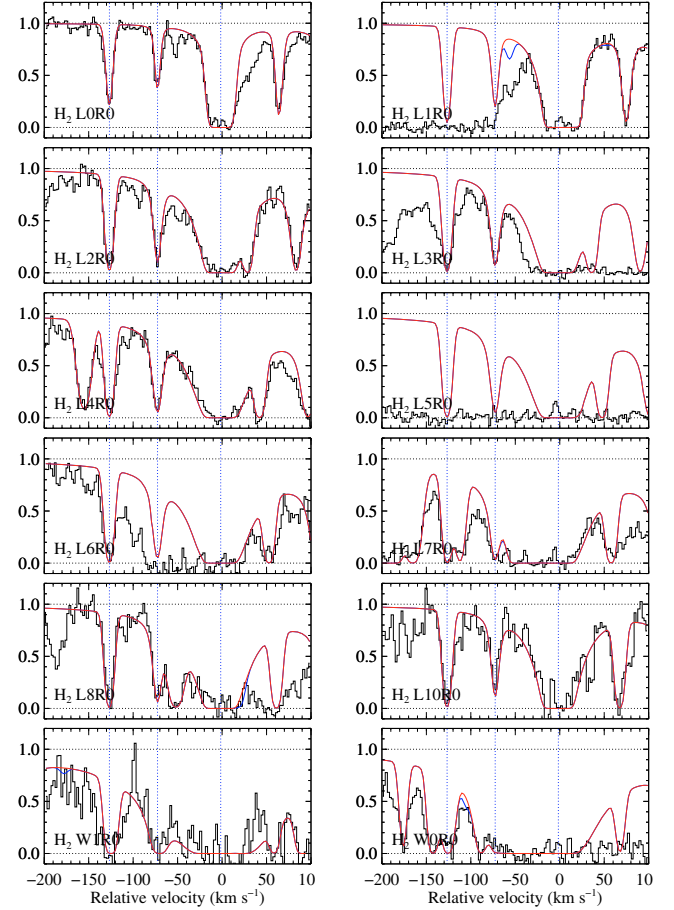


Fig. 16. Fit to $\text{H}_2(J = 0)$ lines. The blue profile is the contribution from HD.

$T_{\text{CMB}}(z = 2.69) = 10.05\text{ K}$. This provides a strong positive test of the hot Big-Bang theory.

Small velocity shifts (of the order of 1 km s^{-1}) are observed between the different neutral and molecular species in the main component. For H_2 , the shift might appear more important but this is rather because the H_2 component is a blend of several sub-components. Neutral chlorine is likely a better indicator of the strongest molecular component (Sonnentrucker et al. 2006). Srianand et al. (2010) recently observed similar small velocity shifts between H_2 and 21-cm absorption in a high redshift DLA which also testifies to the presence of inhomogeneities of the ISM on very small scales. In turn, because the distribution of H^0 and O^0 are expected to follow each other closely, the O I 1302 absorption profile indicates that atomic hydrogen is distributed over the full $\sim 400\text{ km s}^{-1}$ velocity range over which metals are observed. This explains the large velocity shift between the centroid of HI and molecular lines.

All this reinforces the view that the ISM is patchy, with small and dense molecular cloudlets (probably with an onion-like structure) immersed in a warmer diffuse atomic medium (see e.g. Petitjean et al. 1992). We can derive an upper limit on the size of the molecular-rich region along the line-of-sight by considering that all H^0 is associated with the main velocity component of density $n_{\text{H}} = 50\text{ cm}^{-3}$. The corresponding characteristic length is $l = N(\text{H}^0)/n_{\text{H}^0} \approx 0.6\text{ pc}$. A lower-limit can be given using Eq. (A7) of Jura & York (1978; see also Sonnentrucker et al. 2002) and the observed molecular and neutral chlorine fractions. This gives $l > 0.05\text{ pc}$. The molecular region of the system

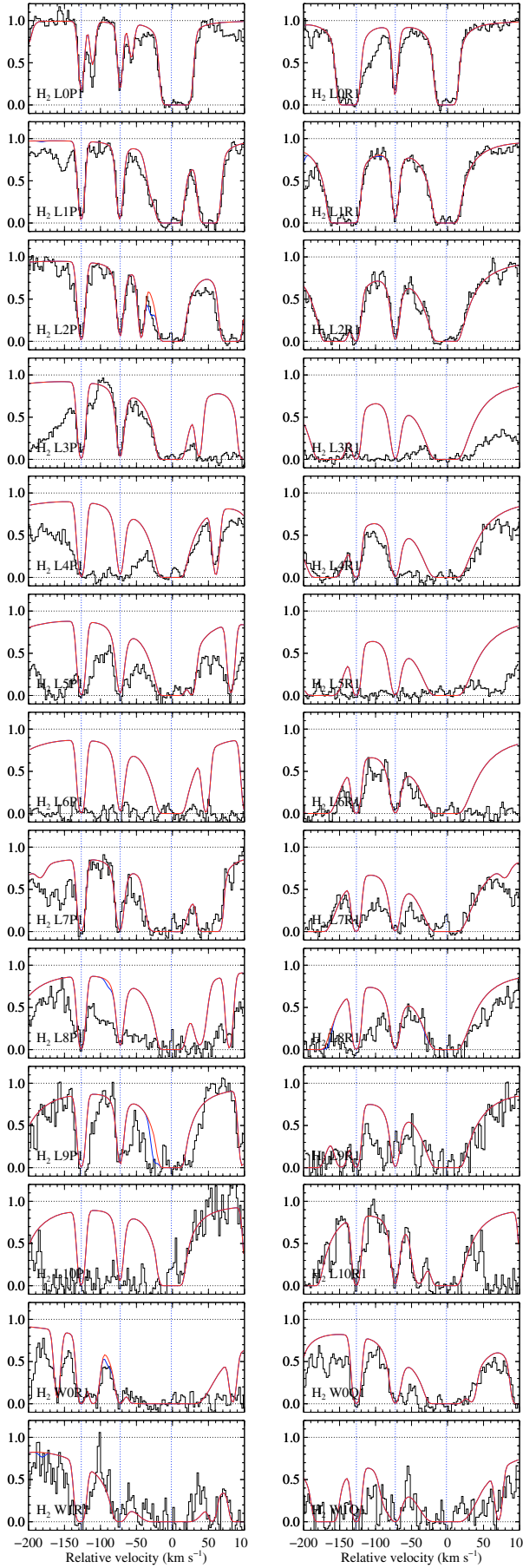


Fig. 17. Fit to $H_2(J = 1)$ lines.

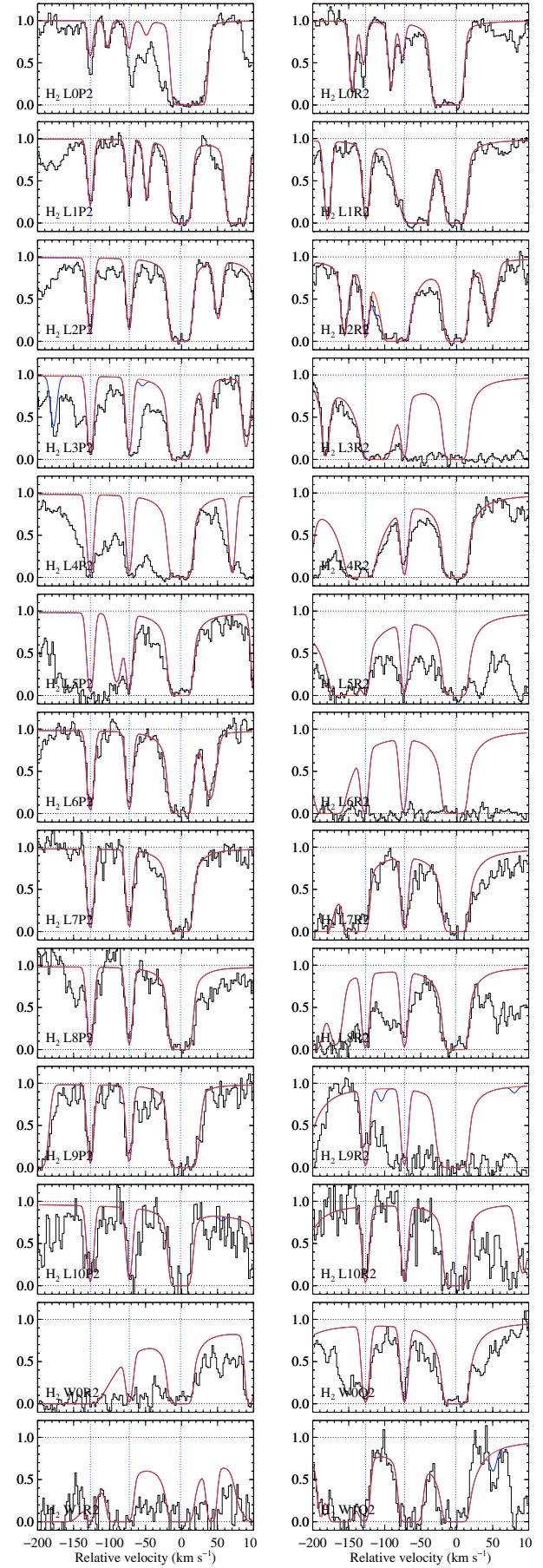


Fig. 18. Fit to $H_2(J = 2)$ lines.

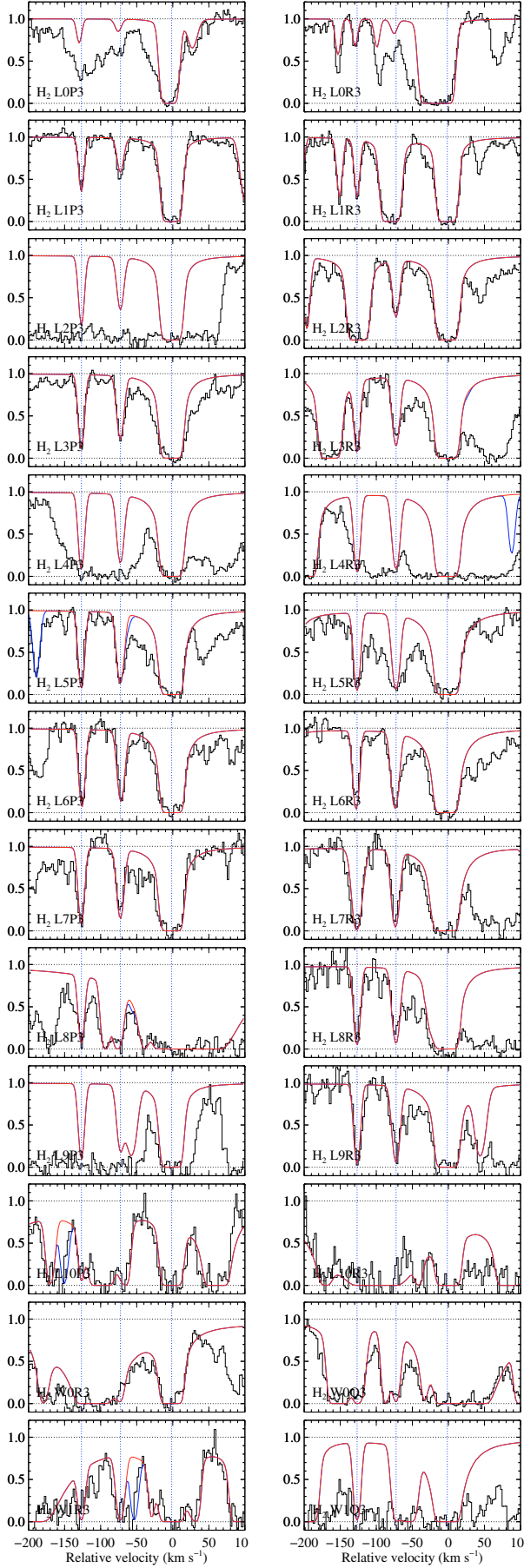


Fig. 19. Fit to $H_2(J = 3)$ lines.

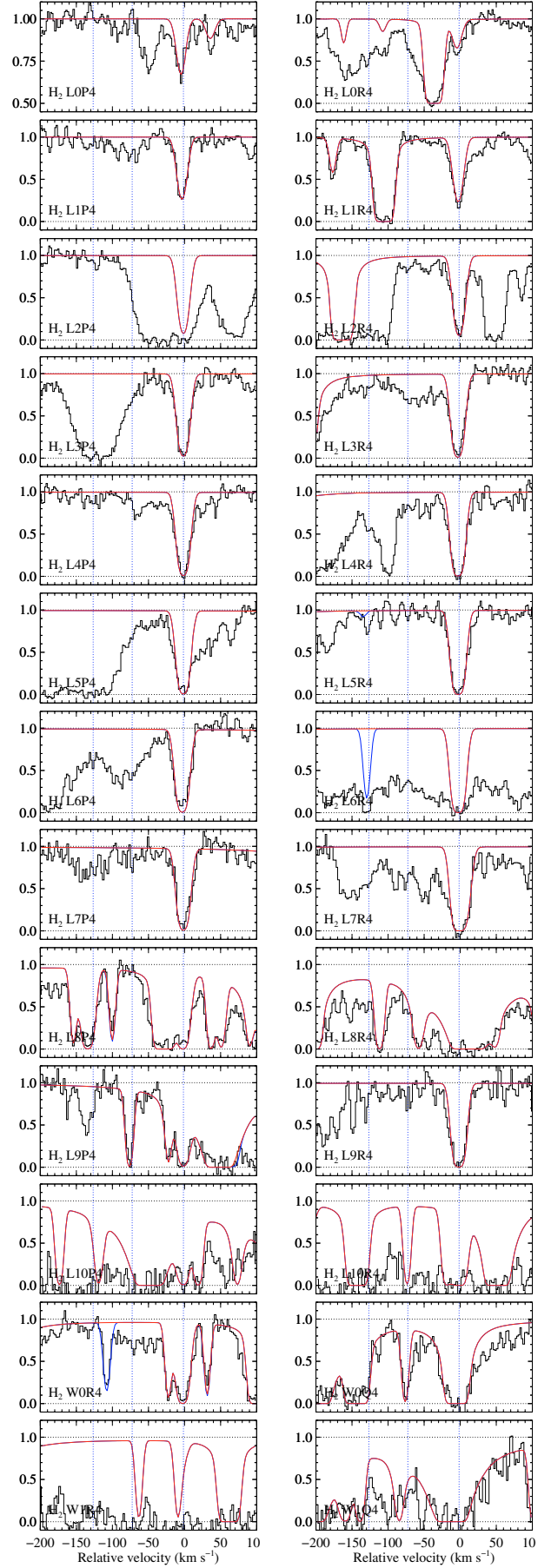


Fig. 20. Fit to $H_2(J = 4)$ lines.

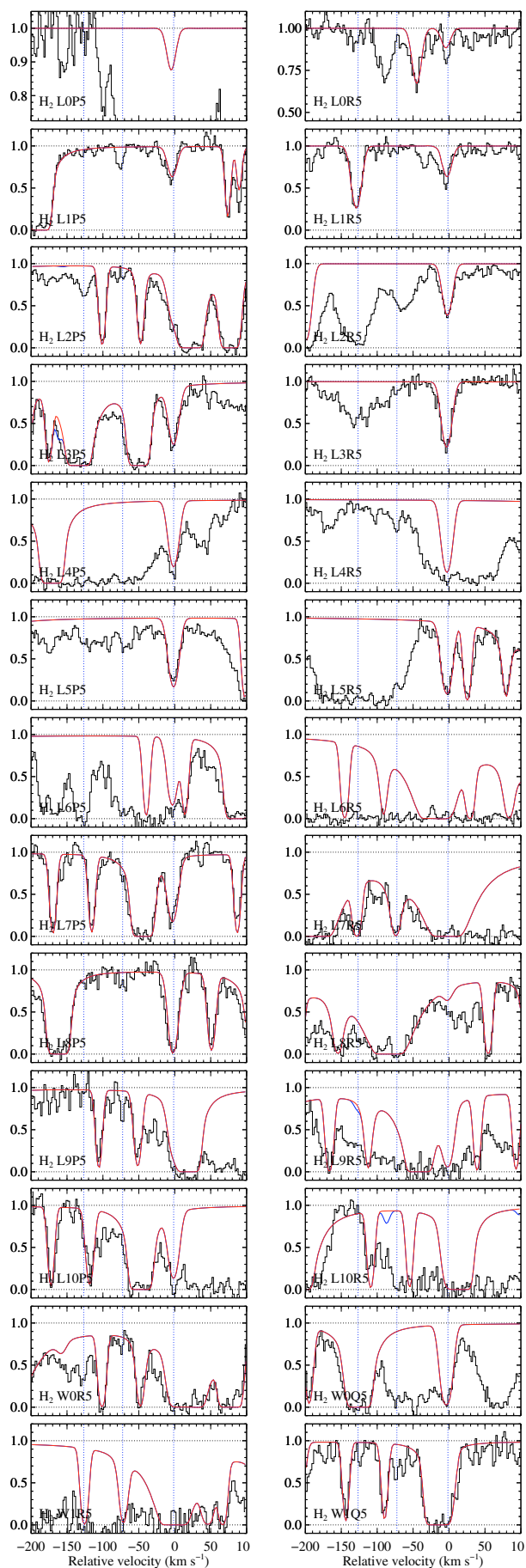


Fig. 21. Fit to $H_2(J = 5)$ lines.

has therefore a very small size and hence small cross-section. It is therefore not surprising that detections of translucent clouds were elusive till now. Studying the frequency of CO absorbers would give an idea of the filling factor of the molecular-rich gas, but requires more extended statistics. Small physical extents could yield partial covering of the background source by the cloud. This may happen in particular if some absorption lines are redshifted on top of emission lines from the extended QSO broad line region, as seen in the case of Q 1232+082 (Ivanchik et al. 2010; Balashev et al. 2010).

Interestingly, the conclusion that the ISM at high redshift is made of small cold cloudlets immersed in warmer diffuse medium has been reached by Gupta et al. (2009) while considering the distribution of cold gas detected in 21-cm absorption. Note that J 1237+0647 is detected in the radio domain by FIRST. However, the low radio flux (2.3 mJy) prevents any spectroscopic study with current radio-telescopes.

Several chemical reactions can take place in this cloud. Indeed, three molecular species are detected, while the presence of neutral chlorine suggests chemical reactions involving HCl and HCl^+ (e.g. Dalgarno et al. 1974; Neufeld & Wolfire 2009). The DLA system toward J 1237+0647 is therefore an excellent candidate to target with future extremely large telescopes (ELTs) to detect other molecular species like C_2 , CH, OH and study astrochemistry in the interstellar medium of high-redshift galaxies. Given the expected attenuation of the quasar by CO-bearing clouds, X-shooter, with its high throughput and medium resolution, is the best instrument to survey carefully selected DLA/sub-DLA candidates for CO absorptions and build a sample of molecular-rich clouds at high redshift, which may then be studied in details in the optical and sub-millimeter ranges with future facilities like ELTs and ALMA.

Acknowledgements. P.N. is supported by a CONICYT/CNRS fellowship and gratefully acknowledges the European Southern Observatory for hospitality during part of the time this work was done. S.L. is supported by FONDECYT grant No. 1100214. We thank Alain Smette for helpful discussions and an anonymous referee for helpful comments and suggestions that improved the content and presentation of the paper. We thank A. Peter for smoothing out the English. We acknowledge the use of the Sloan Digital Sky Survey.

References

- Abgrall, H., & Roueff, E. 2006, A&A, 445, 361
 Abgrall, H., Roueff, E., Launay, F., & Roncin, J.-Y. 1994, Can. J. Phys., 72, 856
 Asplund, M., Grevesse, N., Sauval, A. J., & Scott, P. 2009, ARA&A, 47, 481
 Bailey, D., Salumbides, E. J., Vervloet, M., & Ubachs, W. 2010, Molec. Phys., in press
 Balashev, S. A., Petitjean, P., Ivanchik, A. V., et al. 2010, MNRAS, submitted
 Burgh, E. B., France, K., & McCandliss, S. R. 2007, ApJ, 658, 446
 Burgh, E. B., France, K., & Jenkins, E. B. 2010, ApJ, 708, 334
 Cui, J., Bechtold, J., Ge, J., & Meyer, D. M. 2005, ApJ, 633, 649
 Dalgarno, A., de Jong, T., Oppenheimer, M., & Black, J. H. 1974, ApJ, 192, L37
 Dekker, H., D'Odorico, S., Kaufer, A., Delabre, B., & Kozłowski, H. 2000, in Optical and IR Telescope Instrumentation and Detectors, ed. M. Iye, A. F. Moorwood, Proc. SPIE, 4008, 534
 Eidelsberg, M., & Rostas, F. 2003, ApJS, 145, 89
 Ellison, S. L., Hall, P. B., & Lira, P. 2005, AJ, 130, 1345
 Fontana, A., & Ballester, P. 1995, The Messenger, 80, 37
 Fynbo, J. P. U., Starling, R. L. C., Ledoux, C., et al. 2006, A&A, 451, L47
 Fynbo, J. P. U., Jakobsson, P., Prochaska, J. X., et al. 2009, ApJS, 185, 526
 Goldoni, P., Royer, F., François, P., et al. 2006, in SPIE Conf. Ser., 6269
 Gordon, K. D., Clayton, G. C., Misselt, K. A., Landolt, A. U., & Wolff, M. J. 2003, ApJ, 594, 279
 Gupta, N., Srianand, R., Petitjean, P., Noterdaeme, P., & Saikia, D. J. 2009, MNRAS, 398, 201
 Hirashita, H., & Ferrara, A. 2005, MNRAS, 356, 1529
 Ivanchik, A. V., Petitjean, P., Balashev, S. A., et al. 2010, MNRAS, 404, 1583
 Jenkins, E. B., & Peimbert, A. 1997, ApJ, 477, 265
 Jenkins, E. B., & Tripp, T. M. 2001, ApJS, 137, 297

- Jorgenson, R. A., Wolfe, A. M., Prochaska, J. X., & Carswell, R. F. 2009, *ApJ*, 704, 247
- Jura, M. 1974, *ApJ*, 190, L33
- Jura, M., & York, D. G. 1978, *ApJ*, 219, 861
- Kanekar, N., & Chengalur, J. N. 2003, *A&A*, 399, 857
- Lacour, S., André, M. K., Sonnentrucker, P., et al. 2005a, *A&A*, 430, 967
- Lacour, S., Ziskin, V., Hébrard, G., et al. 2005b, *ApJ*, 627, 251
- Ledoux, C., Petitjean, P., & Srianand, R. 2003, *MNRAS*, 346, 209
- Ledoux, C., Petitjean, P., Fynbo, J. P. U., Møller, P., & Srianand, R. 2006a, *A&A*, 457, 71
- Ledoux, C., Petitjean, P., & Srianand, R. 2006b, *ApJ*, 640, L25
- Ledoux, C., Vreeswijk, P. M., Smette, A., et al. 2009, *A&A*, 506, 661
- Linsky, J. L., Draine, B. T., Moos, H. W., et al. 2006, *ApJ*, 647, 1106
- Markwardt, C. B. 2009, in *Astronomical Data Analysis Software and Systems XVIII*, ASP Conf. Ser., 411, 251
- Morton, D. C. 2003, *ApJS*, 149, 205
- Morton, D. C., & Noreau, L. 1994, *ApJS*, 95, 301
- Neufeld, D. A., & Wolfire, M. G. 2009, *ApJ*, 706, 1594
- Noterdaeme, P., Ledoux, C., Petitjean, P., et al. 2007a, *A&A*, 474, 393
- Noterdaeme, P., Petitjean, P., Srianand, R., Ledoux, C., & Le Petit, F. 2007b, *A&A*, 469, 425
- Noterdaeme, P., Ledoux, C., Petitjean, P., & Srianand, R. 2008a, *A&A*, 481, 327
- Noterdaeme, P., Petitjean, P., Ledoux, C., Srianand, R., & Ivanchik, A. 2008b, *A&A*, 491, 397
- Noterdaeme, P., Ledoux, C., Srianand, R., Petitjean, P., & Lopez, S. 2009a, *A&A*, 503, 765
- Noterdaeme, P., Petitjean, P., Ledoux, C., & Srianand, R. 2009b, *A&A*, 505, 1087
- Péquignot, D., & Aldrovandi, S. M. V. 1986, *A&A*, 161, 169
- Péroux, C., Dessauges-Zavadsky, M., Kim, T., McMahon, R. G., & D'Odorico, S. 2002, *Ap&SS*, 281, 543
- Péroux, C., Dessauges-Zavadsky, M., D'Odorico, S., Kim, T.-S., & McMahon, R. G. 2007, *MNRAS*, 382, 177
- Petitjean, P., Bergeron, J., & Puget, J. L. 1992, *A&A*, 265, 375
- Petitjean, P., Srianand, R., & Ledoux, C. 2000, *A&A*, 364, L26
- Petitjean, P., Srianand, R., & Ledoux, C. 2002, *MNRAS*, 332, 383
- Petitjean, P., Ledoux, C., Noterdaeme, P., & Srianand, R. 2006, *A&A*, 456, L9
- Pettini, M., Zych, B. J., Murphy, M. T., Lewis, A., & Steidel, C. C. 2008, *MNRAS*, 391, 1499
- Prochaska, J. X., Sheffer, Y., Perley, D. A., et al. 2009, *ApJ*, 691, L27
- Prodanović, T., & Fields, B. D. 2008, *J. Cosmol. Astro-Part. Phys.*, 9, 3
- Rachford, B. L., Snow, T. P., Tumlinson, J., et al. 2002, *ApJ*, 577, 221
- Reimers, D., Baade, R., Quast, R., & Levshakov, S. A. 2003, *A&A*, 410, 785
- Savage, B. D., Bohlin, R. C., Drake, J. F., & Budich, W. 1977, *ApJ*, 216, 291
- Sheffer, Y., Federman, S. R., & Andersson, B. 2003, *ApJ*, 597, L29
- Sheffer, Y., Rogers, M., Federman, S. R., et al. 2008, *ApJ*, 687, 1075
- Silva, A. I., & Viegas, S. M. 2002, *MNRAS*, 329, 135
- Snow, T. P., & McCall, B. J. 2006, *ARA&A*, 44, 367
- Sonnentrucker, P., Friedman, S. D., Welty, D. E., York, D. G., & Snow, T. P. 2002, *ApJ*, 576, 241
- Sonnentrucker, P., Friedman, S. D., & York, D. G. 2006, *ApJ*, 650, L115
- Sonnentrucker, P., Welty, D. E., Thorburn, J. A., & York, D. G. 2007, *ApJS*, 168, 58
- Srianand, R., Petitjean, P., & Ledoux, C. 2000, *Nature*, 408, 931
- Srianand, R., Petitjean, P., Ledoux, C., Ferland, G., & Shaw, G. 2005, *MNRAS*, 362, 549
- Srianand, R., Gupta, N., Petitjean, P., Noterdaeme, P., & Saikia, D. J. 2008a, *MNRAS*, 391, L69
- Srianand, R., Noterdaeme, P., Ledoux, C., & Petitjean, P. 2008b, *A&A*, 482, L39
- Srianand, R., Gupta, N., Petitjean, P., Noterdaeme, P., & Ledoux, C. 2010, *MNRAS*, 405, 1888
- Steigman, G. 2007, *Ann. Rev. Nucl. Part. Sci.*, 57, 463
- Tumlinson, J., Prochaska, J. X., Chen, H.-W., Dessauges-Zavadsky, M., & Bloom, J. S. 2007, *ApJ*, 668, 667
- Tumlinson, J., Malec, A. L., Carswell, R. F., et al. 2010, *ApJ*, 718, L156
- van der Tak, F. F. S., Black, J. H., Schöier, F. L., Jansen, D. J., & van Dishoeck, E. F. 2007, *A&A*, 468, 627
- van Dishoeck, E. F., & Black, J. H. 1989, *ApJ*, 340, 273
- Vanden Berk, D. E., Richards, G. T., Bauer, A., et al. 2001, *AJ*, 122, 549
- Vladilo, G., Prochaska, J. X., & Wolfe, A. M. 2008, *A&A*, 478, 701
- Wannier, P., Penprase, B. E., & Andersson, B. 1997, *ApJ*, 487, L165
- Warin, S., Benayoun, J. J., & Viala, Y. P. 1996, *A&A*, 308, 535
- Welty, D. E., Frisch, P. C., Sonneborn, G., & York, D. G. 1999a, *ApJ*, 512, 636
- Welty, D. E., Hobbs, L. M., Lauroesch, J. T., et al. 1999b, *ApJS*, 124, 465
- Wolfe, A. M., Howk, J. C., Gawiser, E., Prochaska, J. X., & Lopez, S. 2004, *ApJ*, 615, 625