



HAL
open science

Jet and torus orientations in high redshift radio galaxies

G. Drouart, C. de Breuck, J. Vernet, R. A. Laing, N. Seymour, D. Stern, M. Haas, E. A. Pier, B. Rocca-Volmerange

► **To cite this version:**

G. Drouart, C. de Breuck, J. Vernet, R. A. Laing, N. Seymour, et al.. Jet and torus orientations in high redshift radio galaxies. *Astronomy & Astrophysics - A&A*, 2012, 548, 10.1051/0004-6361/201220059 . hal-03645711

HAL Id: hal-03645711

<https://hal.science/hal-03645711v1>

Submitted on 6 Jun 2022

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Jet and torus orientations in high redshift radio galaxies[★]

G. Drouart^{1,2,3}, C. De Breuck¹, J. Vernet¹, R. A. Laing¹, N. Seymour³, D. Stern⁴, M. Haas⁵,
E. A. Pier⁶, and B. Rocca-Volmerange²

¹ European Southern Observatory, Karl Schwarzschild Straße 2, 85748 Garching bei München, Germany
e-mail: gdrouart@eso.org

² Institut d'Astrophysique de Paris, 98bis boulevard Arago, 75014 Paris, France

³ CSIRO Astronomy & Space Science, PO Box 76, Epping, NSW 1710, Australia

⁴ Jet Propulsion Laboratory, California Institute of Technology, Mail Stop 169-221, Pasadena, CA 91109, USA

⁵ Astronomisches Institut, Ruhr-Universität Bochum, Universitätsstr. 150, 44801 Bochum, Germany

⁶ Oceanit Laboratories, 828 Fort Street Mall, Suite 600, Honolulu, HI 96813, USA

Received 19 July 2012 / Accepted 19 September 2012

ABSTRACT

We examine the relative orientation of radio jets and dusty tori surrounding the active galactic nucleus (AGN) in powerful radio galaxies at $z > 1$. The radio core dominance $R = P_{\text{core}}^{20\text{GHz}} / P_{\text{extended}}^{500\text{MHz}}$ serves as an orientation indicator, measuring the ratio between the anisotropic Doppler-beamed core emission and the isotropic lobe emission. Assuming a fixed cylindrical geometry for the hot, dusty torus, we derive its inclination i by fitting optically-thick radiative transfer models to spectral energy distributions obtained with the *Spitzer* Space Telescope. We find a highly significant anti-correlation ($p < 0.0001$) between R and i in our sample of 35 type 2 AGN combined with a sample of 18 $z \sim 1$ 3CR sources containing both type 1 and 2 AGN. This analysis provides observational evidence both for the Unified scheme of AGN and for the common assumption that radio jets are in general perpendicular to the plane of the torus. The use of inclinations derived from mid-infrared photometry breaks several degeneracies which have been problematic in earlier analyses. We illustrate this by deriving the core Lorentz factor Γ from the R - i anti-correlation, finding $\Gamma \gtrsim 1.3$.

Key words. galaxies: high-redshift – galaxies: active – radio continuum: galaxies – infrared: galaxies – galaxies: jets – quasars: general

1. Introduction

Radio galaxies are among the most luminous objects in the Universe over the entire electromagnetic spectrum. Their powerful radio emission betrays the presence of a central massive black hole (Blandford & Payne 1982) up to a few billion M_{\odot} in mass (McLure et al. 2006; Nesvadba et al. 2011). Early in the development of the subject, local radio galaxies were associated with massive elliptical galaxies (Matthews et al. 1964). It is now well established that powerful radio sources are also hosted by massive galaxies at higher redshift (e.g. De Breuck et al. 2010), as expected from the bulge-black hole mass relation (Ferrarese & Merritt 2000). Radio galaxies appear to be a singular stage in the evolution of massive galaxies, observed during a peak of activity. This phase presents a unique chance to test models of galaxy formation and the interaction between the AGN and their host galaxies (e.g. Nesvadba et al. 2008).

According to current understanding, an active galactic nucleus (AGN) consists of an accretion disk around a supermassive black hole (SMBH; Rees 1984). About 10% of AGN show radio jets (e.g. Best et al. 2005); these are expected to be aligned with the black hole spin axis. A dusty torus has been hypothesized to explain the observed dichotomy between unobscured (type 1) AGN where the observer can see the region close to the black hole directly and obscured (type 2) AGN with a more edge-on view (for a review, see Antonucci 1993). As a result of this geometry, the AGN emission is anisotropic at most wavelengths. The accretion disk is a powerful source of soft and hard X-rays

which are attenuated when passing through the torus. Assuming a universal torus shape, the dominating factor determining the amount of obscuration along the line of sight is the orientation of the torus with respect to the observer. In particular, soft X-rays are sensitive to the hydrogen column density (N_{H}) which varies from 10^{20} cm^{-2} in type 1 AGN to $>10^{24} \text{ cm}^{-2}$ type 2 AGN (e.g. Ibar & Lira 2007).

In the optical domain, the effects of inclination can explain why the Broad Line Regions (BLR) are observed directly in type 1 AGN, while in type 2's, they are mostly obscured or seen in scattered emission (Antonucci & Miller 1985; Tran et al. 1992). In both types, Narrow Line Regions (NLR) are observed further out from the torus. In the case of type 2's, the torus acts as a coronagraph blocking the much brighter central emission, and allows the NLR to be traced out to $\sim 100 \text{ kpc}$ (Reuland et al. 2003; Villar-Martín et al. 2003). Such detailed studies have shown that the NLR is frequently aligned with the radio jets (McCarthy et al. 1987). Moreover, optical polarimetry of type 2 AGN reveals a continuum polarization angle mostly perpendicular to the radio axis, implying that the light passing through torus opening is scattered by dust clouds along the radio jets (di Serego Alighieri et al. 1989, 1993; Cimatti et al. 1993; Hines 1994; Vernet et al. 2001). Both observations imply that the radio jets are aligned orthogonally to the equatorial plane of the torus.

The torus surrounding the SMBH re-processes a significant fraction of the AGN radiation (X-rays, UV and optical) into mid-IR thermal dust emission. This establishes a radial temperature gradient within the torus ranging from the sublimation temperature ($\sim 1500 \text{ K}$) at the inner surface of the torus to a few

[★] Figures 11, 12, and Tables 1, 2, 6 are available in electronic form at <http://www.aanda.org>

hundred Kelvin in the outer parts. The torus geometry creates a strong anisotropy in the hot dust emission because the amount of extinction towards the innermost (hottest) parts is very sensitive to the orientation with respect to the observer (Pier & Krolik 1992, hereafter PK92). It should therefore be possible to derive the inclination by modelling of the mid-IR SED. The magnitude of the variation with inclination is illustrated by the difference between the mean type 1 and type 2 mid-IR SEDs of an isotropically selected AGN sample where the type 2 AGN SED can be reproduced by simple reddening of the type 1 distribution (Haas et al. 2008).

The most isotropic emission from AGN comes from the radio lobes which mark the interaction between the jets and the surrounding intergalactic medium. This can happen on scales as large as several Mpc and produces steep-spectrum synchrotron emission. In contrast, the radio cores generally have flatter spectral indices and are anisotropic as they are subject to Doppler beaming effects. The radio cores will look brighter if the jet axis is observed closer to the line of sight (e.g. in type 1 AGN). The ratio of the core to total radio emission (core dominance) is a proxy for orientation (e.g. Scheuer & Readhead 1979; Kapahi & Saikia 1982).

In summary, previous observations suggest that the radio jets are orthogonal to the equatorial plane of the torus. Assuming a generic torus geometry, an inclination can be derived by fitting the mid-IR SED; the core dominance can also be used to estimate the inclination of the radio jets. Provided that these assumptions are correct, the two measures of inclination should be consistent. In this paper, we make use of the unique database consisting of six-band mid-IR data for 70 radio galaxies spanning $z = 1$ to $z = 5.2$ (Seymour et al. 2007; De Breuck et al. 2010, S07 and DB10 hereafter). We derive the inclination angle by fitting dust emission from the torus, and compare it with the radio core dominance. We indeed find a significant correlation between these parameters, consistent with the previously mentioned observational statements which sample directions in the plane of the sky. We also use this constraint on the orientation to estimate the pc-scale jet speeds from the core/jet dominance.

This paper is organised as follows. Sections 2 and 3 describe the samples. In Sect. 4, we describe our approach to modelling the mid-IR emission, using both empirical correlations and a torus model for a sub-sample with well-sampled AGN dust emission. We discuss the implications in Sect. 5. Throughout this paper, we adopt the current standard cosmological model ($H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_\Lambda = 0.7$, $\Omega_M = 0.3$).

2. Spitzer high redshift radio galaxy sample

Our *Spitzer* high redshift radio galaxy (SHzRG) sample is selected from a compendium of the most powerful radio galaxies known in 2002. From this parent sample of 225 HzRGs with spectroscopic redshifts, a subset was selected in order to sample the redshift-radio power plane evenly out to the highest redshifts available. The full description of this sample is presented by S07. In short, the HzRG are distributed across $1 < z < 5.2$ and have $P^{3 \text{ GHz}} > 10^{26} \text{ W Hz}^{-1}$, where $P^{3 \text{ GHz}}$ is the total luminosity at a rest-frame frequency of 3 GHz (Table 1 of S07).

2.1. Infrared data

The mid-IR data presented here were obtained with the *Spitzer* Space Telescope (Werner et al. 2004) during Cycles 1 and 4. All galaxies were observed in the four Infrared Array

Camera (IRAC; Fazio et al. 2004) channels (3.6, 4.5, 5.8, 8 μm), the 16 μm peak-up mode of the Infrared Spectrograph (IRS; Houck et al. 2004) and the 24 μm channel of the Multiband Infrared Photometer (MIPS; Rieke et al. 2004). The 24 sources with the lowest expected background emission were also observed with the 70 and 160 μm channels of MIPS. We refer the reader to S07 and DB10 for the full description of the data reduction procedures and the full photometric data (Table 3 of DB10). We augment our photometry with *K*-band magnitudes from S07.

2.2. Radio data and core dominance

The radio morphologies of HzRGs are dominated by steep-spectrum radio lobes with fainter and flatter spectrum radio cores (Carilli et al. 1997; Pentericci et al. 2000). A variety of physical processes contribute to the radio emission at different restframe frequencies (e.g. Blundell et al. 1999). Above $\nu \approx 1 \text{ GHz}$, Doppler boosting will introduce an orientation bias in the radio cores. In addition, the radio lobes are affected by synchrotron and inverse Compton losses. At very low frequencies $\nu \lesssim 100 \text{ MHz}$, synchrotron self-absorption, free-free absorption and any low-energy cutoff of the relativistic particles may reduce the observed synchrotron emission. We therefore choose a rest-frame wavelength of 500 MHz to obtain an accurate measure of the energy injected by the AGN which is also independent of orientation. Observationally, the 500 MHz luminosities $P_{\text{extended}}^{500 \text{ MHz}}$ can also be determined most uniformly over the entire sky (DB10).

Our measure of orientation for the radio jets is the ratio of core to extended emission or core dominance R (Scheuer & Readhead 1979; Kapahi & Saikia 1982). Since the flatter spectrum cores can only be spatially resolved with interferometers at high frequencies, this is defined as $R = P_{\text{core}}^{20 \text{ GHz}} / P_{\text{extended}}^{500 \text{ MHz}}$, where $P_{\text{core}}^{20 \text{ GHz}}$ is the 20 GHz restframe core luminosity.

Core flux densities are either taken from the literature (see Table 1), or measured directly from radio maps by the following method. First, we take the *Spitzer*/IRAC 3.6 μm image and overlay the radio contours. After identifying the core with the host galaxy, we measure the integrated total and core flux densities using the AIPS verb TVSTAT. If only the lobes are identified, we derive an upper limit to the core dominance from the 3σ sensitivity of the radio map. Next, we calculate $P_{\text{core}}^{20 \text{ GHz}}$ using 8.4 GHz observations and the core spectral index α_{core} , defined as $S \propto \nu^\alpha$. If no spectral index is available, we use the median value from our sample $\langle \alpha_{\text{core}} \rangle = -0.8$. Note that given the large range of R , the exact value of α_{core} does not significantly affect our results. At higher redshifts, inverse Compton losses can increase significantly, which may affect R values. However, Fig. 1 does not show a significant dependence of R on redshift, so we consider this effect to be negligible.

Table 1 lists the values of $P_{\text{extended}}^{500 \text{ MHz}}$ (summing the components in the case of multiple detections) and R , together with the core flux densities and spectral indices α_{core} (see references in Table 1).

3. 3CR sample

To compare the results of our approach for type 1 and 2 radio-loud AGN, samples of the two classes with matched selection criteria are required. Unfortunately, there is no type 1 (quasar) sample in the literature which matches our radio-galaxy sample and has both radio and *Spitzer* observations. The best sample available, although significantly smaller than our SHzRG sample

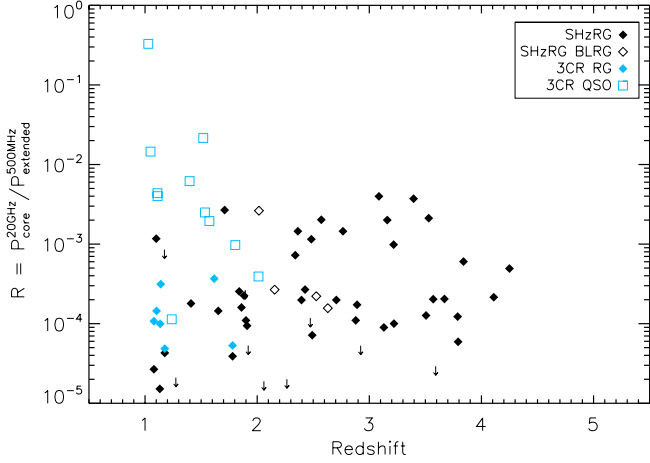


Fig. 1. Core dominance R versus redshift. Note the lower redshifts for the s3CR sample (blue diamonds and open squares). The open black diamonds correspond to radio galaxies with observed broad lines, see Sect. 5.3.

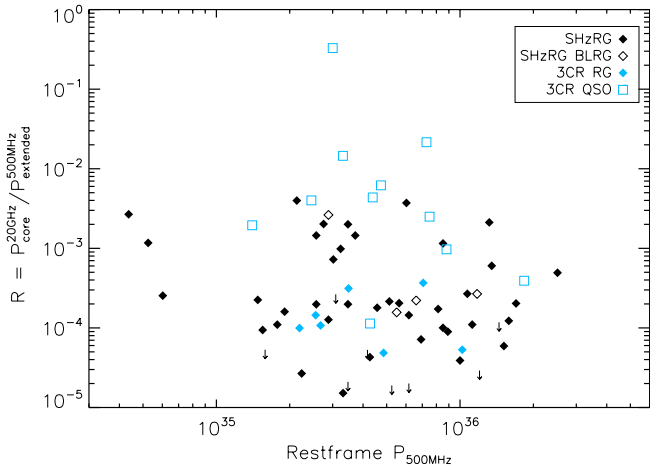


Fig. 2. Core dominance R versus $P_{500\text{MHz}}^{\text{restframe}}$. Note that the s3CR quasars (open blue squares) typically have higher values of R than the galaxies from the same sample (filled blue diamonds).

and with a lower median redshift, is the 3CR sample (Spinrad et al. 1985).

3.1. Infrared data

A selection of 64 3CR high-redshift sources has been observed with the four IRAC channels (3.6, 4.5, 5.8, 8 μm), IRS (16 μm) and the MIPS1 channel (24 μm). We refer to Haas et al. (2008) for a full description of the sample and data reduction. Table 2 reports the *Spitzer* photometry.

3.2. Radio data and core dominance

To compare the 3CR and SHzRG sources, we need equivalent radio data for both groups. We take the relevant measurements for the subset of 3CR sources from Hoekstra et al. (1997). Selecting sources with both *Spitzer* and radio data gives us a subset of 18 sources: 11 quasars and 7 radio galaxies. The radio data are reported in Table 3. We recalculate the 500 MHz rest-frame luminosity from the 178 MHz flux densities (Hoekstra et al. 1997, and references therein), using a spectral index $\alpha = -1.5$, which is typical of powerful steep-spectrum radio sources in

Table 3. Radio data for the s3CR sample.

Name	Type	z	$\log P_{500\text{MHz}}^{\text{restframe}}$ [W Hz $^{-1}$]	$S_{5\text{GHz}}^{\text{core}}$ [mJy]	R
3C 009	Q	2.01	29.26	0.55	0.00039
3C 065*	G	1.18	28.69	0.76	4.9e-05
3C 068.1	Q	1.24	28.63	0.83	0.00011
3C 204	Q	1.11	28.39	0.34	0.0040
3C 205	Q	1.53	28.88	0.67	0.0025
3C 208	Q	1.11	28.64	0.54	0.0044
3C 212	Q	1.05	28.52	0.88	0.015
3C 239*	G	1.78	29.01	0.33	5.3e-05
3C 241	G	1.62	28.85	0.34	0.00037
3C 245	Q	1.03	28.48	1.40	0.33
3C 252	G	1.10	28.41	0.32	0.00014
3C 267	G	1.14	28.54	0.59	0.00031
3C 268.4	Q	1.40	28.68	0.60	0.0062
3C 270.1	Q	1.52	28.86	0.87	0.022
3C 318	Q	1.57	28.15	0.75	0.0019
3C 325	G	1.13	28.34	0.82	9.9e-05
3C 356*	G	1.08	28.43	0.38	0.00011
3C 432	Q	1.80	28.94	0.31	0.00097

Notes. Types: Q = quasar, G = radio galaxy. (*) Radio galaxies common with SHzRG sample.

this frequency range. The 20 GHz rest-frame core luminosities are calculated using the flux density at 5 GHz assuming $\alpha_{\text{core}} = -0.8$. The derived values of R are also reported in Table 3.

4. Modelling the torus emission

4.1. Contributions to the infrared SED

The near- to far-IR emission from radio galaxies and quasars consists of several components: (i) non-thermal synchrotron emission; (ii) line emission; (iii) starburst- and AGN-heated dust continuum; (iv) an (old) stellar population. We now discuss the importance of these contributions in turn.

Of the type 2 objects in our sample, the source which is likely to have the highest synchrotron contribution at 5 μm wavelength is B3 J2330+3927, which has the highest value of R and one of the flattest core spectra ($\alpha = -0.1$). Its extrapolated synchrotron contribution is $<5\%$ of the total emission at 5 μm . The synchrotron contribution can therefore be safely ignored for all of the type 2 objects. However in type 1 objects, this assumption may not be valid. We further discuss this point in Sect. 5.2.2.

The line emission is dominated by fine structure and CO lines in the far-IR (e.g. Smail et al. 2011), and in the mid-IR by polycyclic aromatic hydrocarbon (PAH) emission. The silicate emission/absorption is also seen in type 1/2 AGN, respectively. Both can affect the broad-band photometry. Due to the large redshift range of our sources, the sampling of the SED differs from object to object. To keep our study as homogeneous as possible, we focus only on the part of the SED below the silicate feature ($<10\ \mu\text{m}$), which is dominated by AGN heated dust emission (see below). Mid-IR spectroscopy with IRS (Seymour et al. 2008; Leipski et al. 2010; Rawlings et al., in prep.) shows that PAH features are either not detected or weak relative to the AGN emission.

The contribution of the dust continuum spans three orders of magnitude in wavelength. The energy source of this emission (AGN and/or starburst) has been the subject of considerable debate (e.g. Sanders et al. 1988; Haas et al. 2003). To disentangle these two components, good sampling of the entire IR

SED is essential. The *Herschel* Radio Galaxy Evolution or *pro-jet HeRGÉ* will observe our sample between $70\ \mu\text{m}$ and $500\ \mu\text{m}$. First results on PKS1138–262 (Seymour et al. 2012) show that the AGN component dominates for $\lambda_{\text{rest}} \lesssim 30\ \mu\text{m}$. The contribution of starburst-heated dust may vary from source to source, but is unlikely to contribute significantly at $\lambda_{\text{rest}} \lesssim 10\ \mu\text{m}$ as the required dust temperature would be more than several hundred Kelvin, which cannot be sustained throughout the host galaxy. We consequently focus only on $\lambda_{\text{rest}} \lesssim 10\ \mu\text{m}$. A crucial ingredient in modelling the torus is the $9.7\ \mu\text{m}$ silicate feature, which can be used as an indicator of its intrinsic properties (e.g. the differences between clumpy and continuous or optically thin and thick tori and the dust composition; van Bemmel & Dullemond 2003). Using this feature to characterize the torus is beyond the scope of this paper, since our broad-band photometry is not particularly sensitive to it.

Emission from the old stellar population peaks at $\lambda_{\text{rest}} \sim 1.6\ \mu\text{m}$ (S07). In this paper, we model this contribution assuming a formation redshift $z_{\text{form}} = 10$, using an elliptical-galaxy template predicted by PEGASE.2 (Fioc & Rocca-Volmerange 1997). The stellar contribution from this galaxy template is then normalised for the bluest available band and subtracted in each filter. Uncertainties for the subtracted data are taken to be the rms of the uncertainties of the individual measurement and the stellar fit. For the type 2 3CR radio galaxies, we follow the same method. For the type 1 3CR AGN, we do not subtract any stellar contribution, as the transmitted non-thermal continuum is expected to dominate over the stellar emission. The remaining flux is considered to be a “pure” AGN contribution.

4.2. Sub-samples with well-sampled AGN dust emission

After stellar subtraction (see Sect. 4.1), the signal to noise ratio (S/N) is calculated for each data-point. Then, we create a preferred sub-sample fulfilling the criteria that: (a) there are at least 3 points with $S/N > 2$; and (b) there is no significant contribution from emission with $\lambda_{\text{rest}} > 10\ \mu\text{m}$ in the passband. This in practice restricts the wavelength range to $2\ \mu\text{m} < \lambda_{\text{rest}} < 8\ \mu\text{m}$ in any passband. 35 radio galaxies from the SHzRG sample and 50 sources (22 radio galaxies, 25 quasars and 3 unidentified) from the 3CR sample meet these criteria. We refer to these two restricted samples as “sSHzRG” and “s3CR” respectively. Note that only 31 objects have a known core dominance value in sSHzRG and 18 in the s3CR (11 quasars and 7 radio galaxies).

4.3. Empirical approach

To first order, increasing the viewing angle of the torus causes it to act as a varying dust screen: higher inclinations lead to increased extinction of the hottest dust located in the innermost parts of the torus. Leipski et al. (2010) and Haas et al. (2008) have shown that the mean radio galaxy SED can be approximated by applying an extinction law to the mean quasar SED.

Following the same approach, we model this dust absorption using the mean Sloan Digital Sky Survey (SDSS) quasar spectrum (Richards et al. 2006) and a Fitzpatrick (1999) extinction law with classical Galactic dust properties ($R_V = 3.1$, Fig. 3). This law extrapolates Galactic dust properties to the mid-IR without any specific treatment of the silicate absorption and emission features around $10\ \mu\text{m}$. This latter approximation is still valid as we are focusing on the $2\text{--}8\ \mu\text{m}$ part of the SED.

Using a standard χ^2 minimization technique, we fit two parameters (extinction A_V and normalization) to the data for our

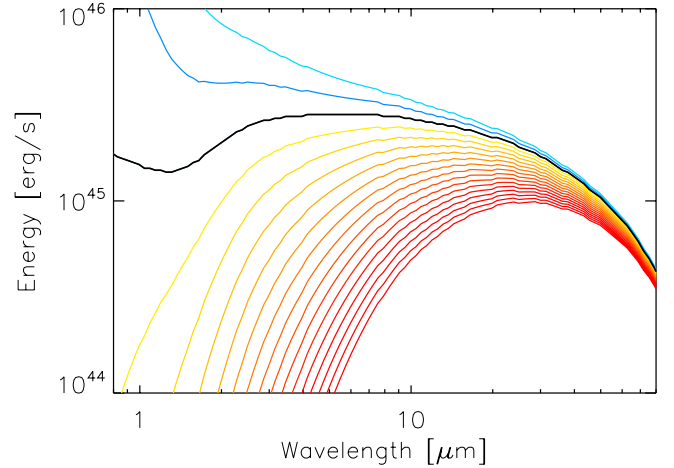


Fig. 3. Template SEDs for the empirical approach. The black curve is the mean quasar spectrum from Richards et al. (2006). The blue/redder curves correspond to a increasing extinction (from -12 to $90 A_V$ in steps of 6), using the Fitzpatrick (1999) extinction law.

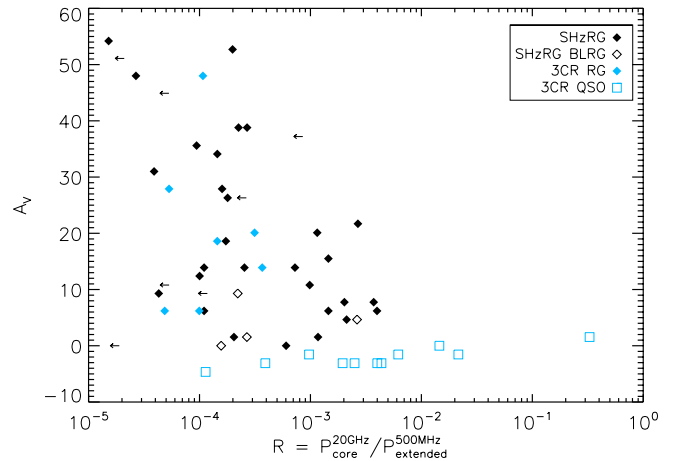


Fig. 4. Extinction A_V versus core dominance R . Note the absence of points in the top right part of the plot. This figure is consistent with Fig. 12 of Cleary et al. (2007).

sSHzRG and s3CR samples. Tables 5 and 6 report the resulting values of A_V . Figure 4 plots A_V against R for the subset of objects with radio data available. Note the apparently unphysical negative A_V values for the majority of s3CR quasars, which are bluer than the composite SDSS template. The lack of points in the upper right part of this plot indicates an absence of core-dominated objects with high extinction, as expected from the orientation-based Unified model (e.g. Antonucci 1993).

To test the correlation between A_V and core dominance taking the upper limits in R into account, we use the survival analysis package within IRAF. The generalized Spearman rank test gives probabilities of non-correlation of $p = 0.009$ and $p = 0.0001$ for sSHzRG and sSHzRG+s3CR samples, respectively. We note that Cleary et al. (2007) report a similar distribution (their Fig. 12), although with a different core dominance definition and using the silicate optical depth, $\tau_{9.7\ \mu\text{m}}$ as a measure of extinction.

4.4. Torus models

We now aim to reproduce the observed range in extinction assuming a torus geometry observed at varying angles to the

line of sight. The infrared torus SED has been modelled extensively, assuming structures with various degrees of complexity. Despite the evidence for a clumpy structure (Krolik & Begelman 1988), the first models to solve the radiative transfer equations used the continuous density approximation (Pier & Krolik 1992). Later, a treatment of the clumpy structure was given by Nenkova et al. (2002). Currently, models with continuous and clumpy structures are available in the literature with a range of geometries and using different computational techniques to solve the radiative transfer equation. Examples include: continuous models (Pier & Krolik 1992; Dullemond & van Bemmell 2005; Granato & Danese 1994; Rowan-Robinson 1995); clumpy models (van Bemmell & Dullemond 2003; Hönig et al. 2006; Schartmann et al. 2005; Fritz et al. 2006; Nenkova et al. 2008) and bi-phased models (Stalevski et al. 2012).

Clumpy models are closer to reality than continuous models, but the latter require fewer free parameters and remain valid if the inter-clump distance is not significantly larger than the clump size. While more sophisticated modelling may be possible for some individual objects, our aim is to extract global trends from as many sources as possible in our sample. Because of the small number and irregular sampling of the data points in the mid-IR, we therefore opt for the Pier & Krolik continuous model (PK model, hereafter). Trying to derive detailed information on the torus itself is beyond the scope of this paper, and remains challenging even at lower redshift (e.g. Ramos Almeida et al. 2009; Kishimoto et al. 2011).

The PK model assumes a cylindrical geometry for the torus. The opening angle of the torus is $\theta = \tan^{-1}(2a/h)$, where a/h is the aspect ratio (Fig. 5). It models the spectrum of radiation from the central point sources as a power law decreasing from UV to IR. As the index of this power-law α does not have a significant effect for our purpose, we take the default value $\alpha = 1$. We adopt an effective temperature for the inner edge of the torus of $T_{\text{eff}} = 1000$ K as this is the only temperature available for all the geometries in the template library from PK92 (see Table 4). Indeed, a higher T_{eff} corresponds to a higher contribution from the hottest dust (i.e. 1–2 μm) but would not affect the global analysis presented here.

The inner radius a (see Fig. 5) is then set by the central source luminosity. For each inclination i , an SED is computed by solving the radiative transfer equations. The final SED is the sum of the thermal torus emission and the absorbed quasar component. For a full description of the model and an example of its application, we refer the reader to Pier & Krolik (1992, 1993).

The main goal of this approach is to find a physically meaningful geometry which can give an adequate description of the entire sample, to estimate torus inclinations and thereby to test the general configuration mentioned in Sect. 1. We have considered a set of representative candidate models, as listed in Table 4.

Some reasonable arguments and previous observations help us to define the average geometry and to select among the models in Table 4. First, the statistical study of the relative frequency of type 1 and 2 AGN by Barthel (1989) implies a torus opening angle $\theta \sim 45^\circ$, which corresponds to $a/h \sim 0.3$ in the PK models. Second, X-ray observations of nearby AGN (e.g. NGC 1068) set a lower limit on the torus opacity of $\tau \geq 1$ (Mulchaey et al. 1992). Only the c and w model geometries satisfy both criteria. The main difference between these geometries is illustrated in Fig. 6: they have chunky (c) and disky (w) shapes, respectively.

A more objective approach is to give each model a score based on the goodness of fit averaged over the whole sample. The fit again uses χ^2 minimization with two free parameters per fit: the normalization of each model and the inclination i of the

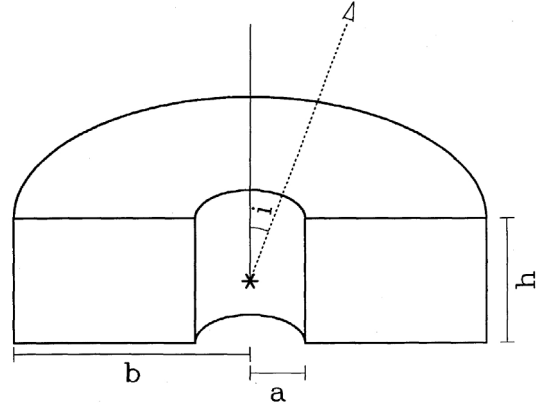


Fig. 5. Sketch of the PK model (Pier & Krolik 1992). The inner radius, a , outer radius, b and height, h are marked, and the torus is viewed from angle i , measured as shown. The central nucleus is indicated by an asterisk. The density is assumed to be constant throughout the torus.

Table 4. Parameters of the torus model.

Model	Parameters				Score
	τ_r	τ_z	a/h	b/h	
a	0.1	0.1	0.3	1.3	59
b	1.0	0.1	0.1	10.1	70
c	1.0	1.0	0.3	1.3	64
d	1.0	1.0	0.2	1.2	68
e	1.0	1.0	0.1	1.1	54
f	0.1	0.1	0.1	1.1	67
gg	1.0	1.0	1.0	2.0	33
hh	0.1	0.1	1.0	2.0	49
ii	10.0	10.0	1.0	2.0	2
j	10.0	1.0	0.1	10.1	50
o	1.0	0.1	1.0	11.0	28
p	10.0	1.0	1.0	11.0	14
t	1.0	0.1	0.3	10.3	59
w	10.0	1.0	0.3	10.3	44
y	10.0	10.0	1.0	2.0	2

Notes. τ_r : radial depth, τ_z vertical depth, a/h ratio of inner radius over height and b/h : outer radius. For the score value, see Sect. 4.4.

torus. This score (rightmost column of Table 4) is the number of “reasonable” fits, i.e. those with an associated probability of $>95\%$. The average score is therefore an indicator of how well the model fits the samples: higher scores correspond to models which performed better at modelling the observed SEDs. Fig. 7 shows the inclination distribution for the four geometries with the best scores for both samples (s3CR and sSHzRG). Models with narrow opening angles (e.g. b and f , and d to a lesser extent) artificially bias the distribution towards low inclinations. We therefore do not use the scores in Table 4 to simply select between all the torus geometries; instead, the scores are used to select only between models c and w , which already meet criteria set by the physical arguments and observations described above. Of these, we select c , which has the higher score (Table 4). The inclinations i from model c with their 5 μm rest frame monochromatic energies of the torus model $\nu P_{\nu(5\mu\text{m})}^{\text{AGN}}$ are reported in Table 5 and Table 6 for the sSHzRG and s3CR samples, respectively. Their SEDs are presented in Figs. 11 and 12. Section 5 discusses these inclinations in more detail.

Figure 8 shows a plot of core dominance R against the inclination i derived from the mid-IR observations. The distribution of the points in the diagram suggests an anti-correlation

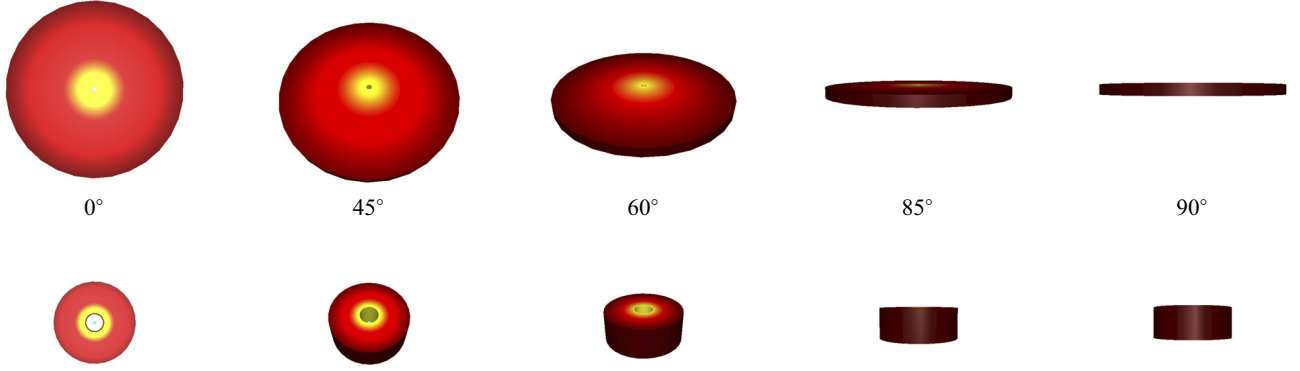


Fig. 6. Modelled tori for $\tau_r/\tau_z = 10$ (top, model *w*, disky) and $\tau_r/\tau_z = 1$ (bottom, model *c*, chunky) for increasing inclinations. Note the appearance of the central engine at 45° (light blue point), and the complete disappearance of the hottest part for 90° (yellow area) due to the cylindrical configuration. This artificially emphasises the difference of the SEDs in the range $85^\circ < i < 90^\circ$.

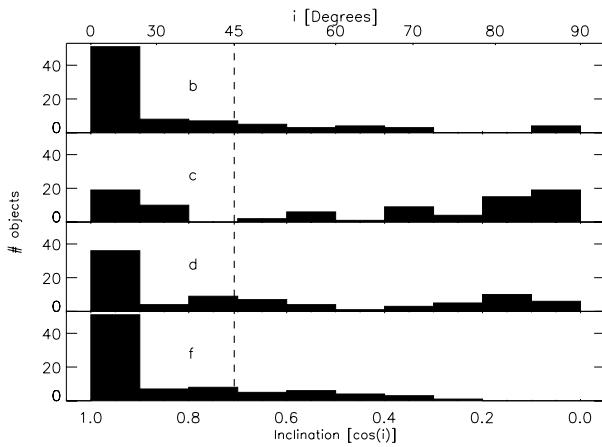


Fig. 7. Distribution of inclination angles i for the four best-scoring models. The vertical dashed line indicates $i = 45^\circ$.

between i and R , which is confirmed with a generalized Spearman rank test for the sSHzRG sample alone ($p = 0.011$). Adding the s3CR sample causes the correlation to become extremely significant ($p < 0.0001$), but note that this is primarily because of the inclusion of the (less obscured) s3CR quasars, which extends the ranges of both R and i .

We would like to add a cautionary note on using the individual inclinations at face value. One source of uncertainty stems from the misalignment of the radio jets with respect to the torus axis. This misalignment adds an intrinsic scatter in the R - i relation. A similar scatter is observed in the plane of the sky through polarimetric measurements.

The distribution of inclinations within $0^\circ < i \lesssim 45^\circ$ is inconsistent with the expectations for an isotropically-distributed parent sample: far too many values are clustered around $i \approx 0^\circ$ or $i \approx 30^\circ$. In particular, inclinations close to $i \approx 0$ seem suspicious since they would be expected for blazars which are absent from our sample.

5. Discussion

5.1. From extinction to inclination

Both the empirical approach (Sect. 4.3) and the torus model (Sect. 4.4) successfully fit the hot dust emission. We now check the consistency between these two methods by comparing the empirical extinction and the orientation from our best torus model. As shown in Fig. 9, these two parameters are tightly

Table 5. Results from the modelling of the SHzRG sample from Sects. 4.3 and 4.4.

Name	A_V	i [$^\circ$]	$\log \nu P_{\nu(5 \mu\text{m})}^{\text{AGN}}$ [W]
6C 0032+412	1.5	46	38.87
MRC 0037-258	1.5	67	38.11
6C 0058+495	37.2	86	38.04
MRC 0114-211	26.3	86	38.40
MRC 0152-209	10.8	79	38.85
MRC 0156-252	4.7	57	39.05
MRC 0211-256	3.1	55	37.92
TXS 0211-122	13.9	79	39.13
MRC 0324-228	26.3	86	38.32
MRC 0350-279	6.2	59	37.99
MRC 0406-244	38.8	86	38.74
PKS 0529-549	34.1	86	38.59
USS 0828+193	7.8	67	39.31
B2 0902+34	7.8	67	38.61
6CE 0905+3955	38.8	86	38.64
3C 239	31.0	86	38.43
MRC 1017-220	4.7	53	38.47
3C 257	9.3	71	39.04
WN J1123+3141	10.8	71	39.27
PKS 1138-262	1.5	67	39.32
6C 1232+39	12.4	75	38.70
4C 24.28	13.9	82	38.90
USS 1410-001	6.2	64	38.83
USS 1558-003	9.3	71	38.91
LBDS 53W002	12.4	79	38.52
3C 356.0	48.0	86	38.07
3C 368	54.2	86	37.91
7C 1805+6332	13.9	79	38.06
TXS J1908+7220	4.7	59	39.43
TN J2007-1316	0.0	31	38.83
MRC 2025-218	0.0	46	38.25
4C 23.56	20.1	82	39.34
MRC 2224-273	4.7	79	38.30
B3 J2330+3927	6.2	75	39.35
3C 470	34.1	86	38.49

correlated. Their relation is well described by the following function:

$$A_V = 8.43(\cos i)^{-0.62} - 10.89. \quad (1)$$

Since the link between extinction and inclination is essentially determined by the geometry of the torus, the small scatter provides support for our choice of a single torus model (*c*) for the entire sample (sSHzRG and s3CR).

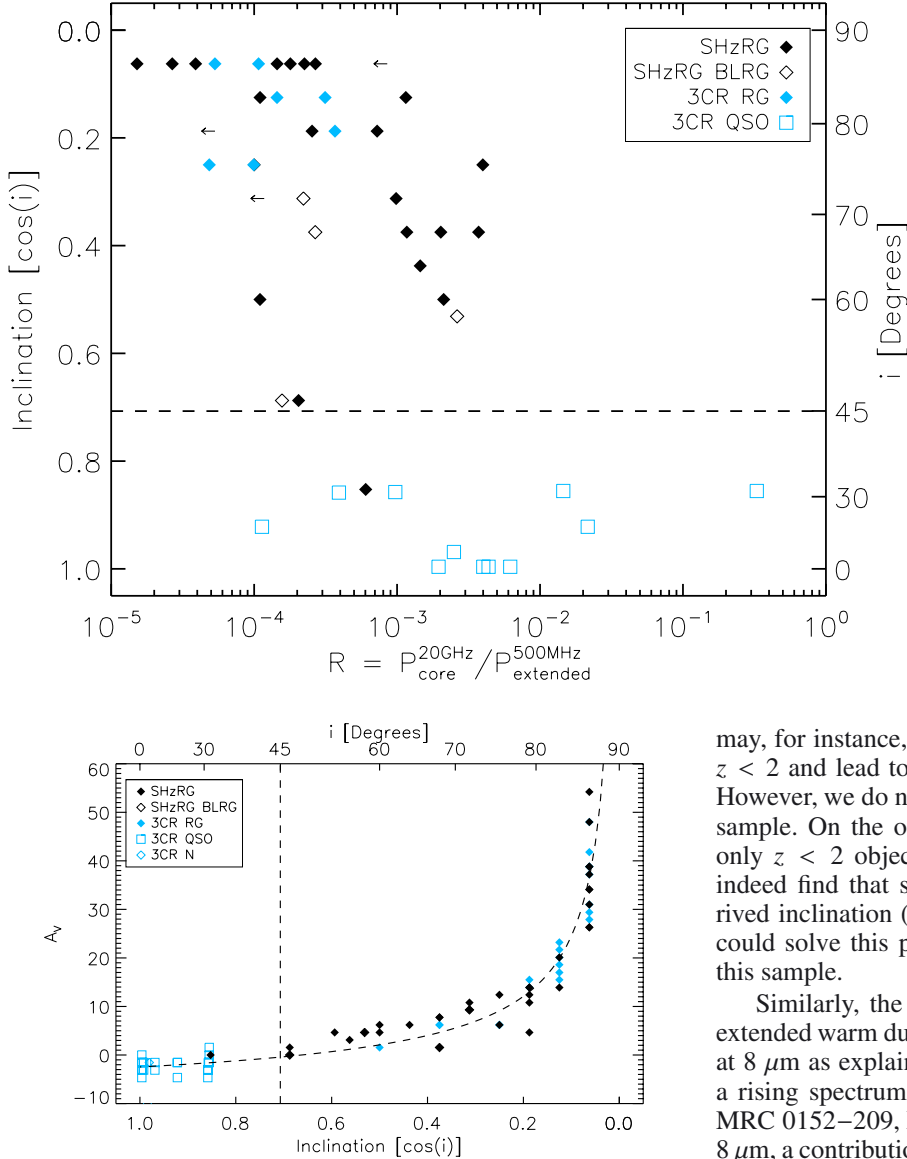


Fig. 8. Plot of inclination i against core dominance R . The dashed line at $i = 45^\circ$ corresponds to the division between radio galaxies and quasars derived by Barthel (1989). Filled diamonds represent the type 2 radio galaxies from the sSHzRG (black) and s3CR (blue) samples, respectively. Open symbols represent: sSHzRG broad-line radio galaxies (black diamonds) and s3CR quasars (blue squares).

Fig. 9. Inclination i obtained from the torus model plotted against extinction A_V from the empirical approach. The dashed vertical line is plotted at $i = 45^\circ$, and the dashed curve corresponds to Eq. (1).

5.2. Isolating the torus emission

5.2.1. Stellar and extended warm dust emission

Here, we investigate the impact of the assumptions made in Sect. 4.1 to isolate the hot dust component from all other contributions in the SED. As we assume that the main stellar contribution is dominated by an evolved population (>500 Myr old), all galaxies of our sample have already formed the bulk of their stellar mass (see S07). The SED of such a population has a characteristic shape for $\lambda_{\text{rest}} \gtrsim 1.6 \mu\text{m}$, as illustrated by the stellar SEDs plotted in Fig. 11. This shape is relatively independent of age, since the light is dominated by low-mass stars. The assumption of a single high formation redshift therefore has minor consequences for the separation of stellar and hot dust components.

A more important aspect is the SED coverage, which varies systematically with redshift. The number of data points measuring purely stellar emission depends on the sampling beyond the stellar bump ($\sim 1.6 \mu\text{m}$) and the relative hot dust contribution in the redmost stellar dominated IRAC channel. This

may, for instance, cause an overestimate of stellar emission for $z < 2$ and lead to a higher inclination and extinction estimate. However, we do not observe such a redshift bias in the sSHzRG sample. On the other hand, the s3CR sample, which contains only $z < 2$ objects, may be more affected by this effect. We indeed find that s3CR type 2 sources mostly have a high derived inclination ($70^\circ < i < 90^\circ$). Adding near-IR photometry could solve this problem, but this is not available for most of this sample.

Similarly, the reddest end of the SED may be affected by extended warm dust emission. To minimize this problem, we cut at $8 \mu\text{m}$ as explained in Sect. 4.1. Nevertheless, in some cases, a rising spectrum from $\lambda_{\text{obs}} = 16 \mu\text{m}$ to $\lambda_{\text{obs}} \geq 24 \mu\text{m}$ (e.g. MRC 0152–209, MRC 0406–244) suggests that even at $\lambda_{\text{rest}} = 8 \mu\text{m}$, a contribution from such a component cannot be excluded.

Indeed, we note that most of the galaxies for which the fit converged to the highest inclination (i.e. the reddest model) may be affected (MRC 0211–256, MRC 0324–228, 6CE 0905+3955, 3C 356, 3C 368, 7C 1805+6332, 3C 470). This could partially explain the clustering of points at $i = 86^\circ$. To test the influence of this contamination, we reduced by a factor of three the flux in the redmost filter for these seven sources. The effect is small: the inclination only decreases in three of them, with a maximum change of 7° . To quantify this effect, a tightly sampled SED over the wavelength range $8 \mu\text{m} < \lambda_{\text{rest}} \leq 20 \mu\text{m}$ would be required. Alternatively, we would have to isolate the torus spatially from more extended emission. This has been attempted for nearby type 2 AGN (van der Wolk et al. 2010) where such an extended component has indeed been identified at $\lambda_{\text{rest}} = 12 \mu\text{m}$. Both types of observation are beyond the reach of current facilities for our samples.

5.2.2. Additional complications for type 1 AGN

In addition to the difficulty of estimating the inclination of an individual type 1 AGN, as described in Sects. 4.3 and 4.4, there are two further complications: at the blue end of the spectrum, transmitted quasar continuum emission can still outshine both the host galaxy and the hot dust emission from the torus. The

PK model already includes the power-law emission from the AGN. As explained in Sect. 4.1, we have not included a contribution from the host galaxy in fitting type 1 AGN. As a test, we have included this component in the same way as in the type 2 AGN, i.e. assuming that the bluest IRAC point is 100% stellar emission. In all 11 s3CR quasars, we find that adding this stellar component severely degrades the quality of the fit, but even in this extreme case, the R - i correlation still remains significant at the $p = 0.002$ level. It is still possible that a smaller stellar component has a smaller contribution on the blue end. Subtracting such a partial stellar population would slightly increase the inclinations found.

On the red end of the torus emission, some sources may also have a contribution from core synchrotron emission (Cleary et al. 2007). This effect is similar to the extended dust component described above, but given the relative small core dominance in our samples (Tables 1 and 3), we do not expect such contributions to have a significant effect.

Despite the above complications, we do find that all quasars have $i < 45^\circ$ (see Fig. 8) as expected from unified models.

5.3. The relative orientations of jets and tori

Our mid-IR SED fitting provides the first estimates of AGN torus inclinations at high redshift. The significant correlation between the radio core dominance and the torus inclination (Fig. 8) implies that the radio jets are indeed generally orthogonal to the equatorial plane of the torus. Orthogonality in projection on the plane of the sky had already been inferred from polarimetric measurements (e.g. Vernet et al. 2001). Our result confirms this geometry in an orthogonal plane containing the line of sight. Together, these observations provide further evidence in favour of the orientation-based unified scheme for AGN.

In the choice of our geometrical model of the torus, we have assumed a half opening angle $\theta = 45^\circ$, as constrained by statistical studies of the relative numbers of type 1 and 2 AGN in a radio-selected sample (Barthel 1989). Other studies have suggested slightly different values of θ (e.g. Willott et al. 2000, find $\theta = 56^\circ$). The torus opening angle may well vary from object to object. This will produce a natural scatter in the R - i relationship. In addition, the transition between direct and obscured view is not expected to be sharp due to the clumpiness of the obscuring material. Two classes of radio galaxies are expected to have inclinations near this transition region. First, radio galaxies with observed broad H α emission (Humphrey et al. 2008; Nesvadba et al. 2011), are indeed found at $45^\circ < i < 75^\circ$ (open diamonds in Fig. 8). Second, the two radio galaxies with observed broad absorption lines (MRC 2025–218 and TXS J1908+7220; Humphrey et al. 2008; De Breuck et al. 2001), have $i = 46^\circ$ and $i = 59^\circ$, respectively, consistent with the expectation to observe them near the grazing line of sight along the torus (e.g. Ogle et al. 1999).

In addition, the torus model adopted to be most representative for the complete sample may not be the optimal choice for a given object. This is particularly true for type 1 AGN, where we cannot constrain the orientation within the range where we have a direct view of the central source, i.e. $0^\circ < i \lesssim 45^\circ$ (see Sects. 4.3, 4.4 and 5.2.2).

Given these caveats, we deliberately do not quote uncertainties on i . In Sect. 5.4 (below), we use only the inclinations for the Type 2 galaxies.

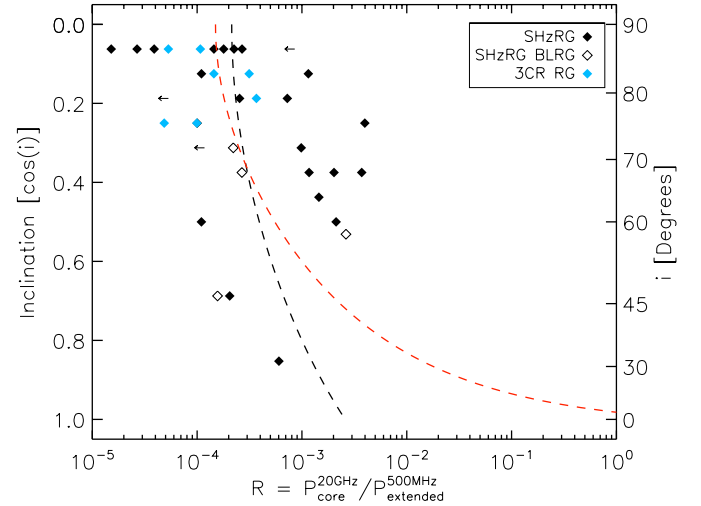


Fig. 10. Core dominance R versus inclination i . The black dashed line is the best fit for all type 2 radio galaxies, $\Gamma = 1.3$. The red dashed line represents the case $\Gamma = 5$.

5.4. Constraints on the core Lorentz factor

Many lines of evidence lead to the conclusion that the jets in radio-loud AGN are relativistic on pc scales, with Lorentz factors of at least 2 and perhaps as high as 50. The highly significant correlation between R and i (Fig. 8) arises naturally if the core radio emission comes from the base of a relativistic jet. If the intrinsic ratio between core and extended emission is constant, then we can use the core dominance for simple assumptions about the jet flow. For anti-parallel, identical jets with velocity βc (Lorentz factor $\Gamma = (1 - \beta^2)^{-1/2}$) and spectral index α at an angle i to the line of sight, the predicted value of R is:

$$R = R_0 \left[(1 - \beta \cos i)^{-(2-\alpha)} + (1 + \beta \cos i)^{-(2-\alpha)} \right]$$

where $2R_0$ is the value of R when the jets are in the plane of the sky ($i = 90^\circ$).

Figure 10 plots R against i for all of the type 2 radio galaxies (sSHzRG and s3CR) with measured R values. For a spectral index $\alpha = -0.9$ (the median value of both samples), the best-fitting Lorentz factor derived from an unweighted least-squares fit to the relation between $\log R$ and i is $\Gamma = 1.3$, with $R_0 = 10^{-4}$. As our type 2 samples are restricted to inclinations $i > 30^\circ$, the Lorentz factor is not well constrained (the dependence of R on i is quite flat). Only very low values of Γ are firmly excluded, and adequate fits can be found for any $\Gamma \gtrsim 1.3$ (for example, see Fig. 10, where the red dashed line represents $\Gamma = 5$). The dispersion in $\log R$ for the best-fitting Γ is 0.59. A stronger constraint on Γ could be obtained from the R - i relation for a comparable sample of broad-lined objects (or just from the median value of R if inclinations are not available). As noted earlier (Sect. 5.3), we do not believe that our estimates of inclination for the s3CR quasars are reliable enough for this purpose.

One potential source of additional scatter in the relation between core dominance and inclination is any dependence of the intrinsic core/extended ratio on luminosity. Such a dependence was found by Giovannini et al. (1988) for a sample of sources covering a wide range in luminosity. We have checked for this effect over the restricted luminosity range of the present type 2 sample, but find no detectable correlation between core dominance and luminosity (Fig. 2).

The most appropriate comparison study in the literature is a Bayesian analysis of the core dominance distribution for a sample of powerful FR II radio sources including both broad and narrow-line objects at $z < 1$ (Mullin & Hardcastle 2009). This gave $\Gamma = 10^{+3}_{-7}$ with a dispersion of 0.62 in $\log R$ assuming a single value of Γ and a dispersion in intrinsic core dominance. A value of $\Gamma = 10$, would also be fully consistent with the present data. The jet/counter-jet ratio on kpc scales provides an independent estimate of orientation. This can be measured adequately only for nearby, low-luminosity radio galaxies, for which Laing et al. (1999) found $\Gamma = 2.4$ and a dispersion of 0.45 in $\log R$.

Various lines of argument suggest that parsec-scale jets have velocity gradients and that the single-velocity model used here is oversimplified. If this is the case, then our analysis is sensitive to the *slowest-moving* component which makes a significant contribution to the rest-frame emission. Faster material will dominate only at smaller angles to the line of sight.

We conclude that the relation between core dominance and inclination for the present sample of high-redshift radio galaxies is fully consistent with their cores being the bases of relativistic jets, as inferred for radio-loud AGN at lower redshift, but that we can set only a lower bound on the Lorentz factor, $\Gamma \gtrsim 1.3$ without additional data.

6. Summary and conclusions

We have examined the relative inclinations of jets and tori in powerful radio-loud AGN at $z > 1$. To estimate the orientation of the radio jets, we have introduced a new definition of the radio core dominance R as the ratio between high-frequency, anisotropic core emission and low-frequency, isotropic extended emission, both measured in the rest frame. To estimate the orientation of the dusty torus, we fit optically-thick radiative transfer models to existing and new *Spitzer* 3.6–24 μm photometry. To isolate the hot dust emission in the torus from non-thermal and host galaxy contributions, we subtract an evolved stellar population in the type 2 sources, and restrict to the $2 \mu\text{m} < \lambda_{\text{rest}} < 8 \mu\text{m}$ region.

Using this method, we derive radio core dominance and torus inclination values for 42 type 2 and 11 type 1 AGN. This mid-IR determination of the inclination allows us to draw the following conclusions:

1. The significant correlation between R and i implies that radio jets are indeed approximately orthogonal to the equatorial plane of the torus as predicted by the orientation-based AGN unified scheme. A similar result in the plane of the sky has been reported using polarimetric measurements, but our results provide additional evidence in the complementary direction, i.e. along the line of sight.
2. The R – i correlation is consistent with the radio cores being the bases of relativistic jets with Lorentz factors $\Gamma \gtrsim 1.3$.
3. Assuming our torus geometry is representative, we can estimate inclinations for larger samples of type 2 AGN using the relation (Eq. (1)) between inclination and the obscuration A_V as derived from a simple reddening of a mean type 1 template.

The present study characterises the anisotropic component of the dust IR SED in type 2 AGN. The remaining more isotropic dust emission, dominating at longer wavelengths, is powered by a combination of AGN and starbursts. Our new *Projet HeRGÉ* aims to disentangle these two contributions by adding *Herschel* 70–500 μm photometry (Seymour et al. 2012) to our sample of 71 radio galaxies at $1 < z < 5.2$.

Acknowledgements. We thank the referee for the careful reading of the manuscript and their constructive comments. The authors thank K. Blundell for useful comments and her participation in the improvement of this paper and C. Leipski for reanalysing some of the 3CR photometry. Nick Seymour is the recipient of an Australian Research Council Future Fellowship. This work is based on observations made with the *Spitzer* Space Telescope, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under contract with NASA.

Appendix A: Notes on individual sources

B3 J2330+3927 This galaxy has the highest core dominance in the SHzRG sample ($R = 0.004$). The radio emission is complex (Pérez-Torres & De Breuck 2005).

6C 0140+326 This source (the second highest redshift source in the SHzRG sample) has been removed from this study, as a foreground object contaminates the galaxy image in the IRAC bands.

4C 60.07, 3C 356, MRC 2048–272, 7C 1756+5620. These four objects have double components in the IRAC maps, but only one of each pair coincides with a radio core.

MRC 2025–218, MRC 0156–252, MRC 1017–220, TXS 1113–178, MRC 1138–262, MRC 1558–003, MRC 0251–273 These galaxies show broad permitted lines (Nesvadba et al. 2011; Humphrey et al. 2008).

6C 0032+412 This galaxy exhibits a very hot dust component in its mid-IR SED (De Breuck et al. 2010).

TNJ2007–1316 The previously quoted flux density at 5.6 μm has been replaced with a 3σ upper limit of $F_{5.6 \mu\text{m}} < 146.0 \mu\text{Jy}$.

References

- Antonucci, R. 1993, ARA&A, 31, 473
 Antonucci, R. R. J., & Miller, J. S. 1985, ApJ, 297, 621
 Athreya, R. M., Kapahi, V. K., McCarthy, P. J., & van Breugel, W. 1997, MNRAS, 289, 525
 Barthel, P. D. 1989, ApJ, 336, 606
 Best, P. N., Longair, M. S., & Roettgering, J. H. A. 1997, MNRAS, 292, 758
 Best, P. N., Kauffmann, G., Heckman, T. M., et al. 2005, MNRAS, 362, 25
 Blandford, R. D., & Payne, D. G. 1982, MNRAS, 199, 883
 Blundell, K. M., Rawlings, S., Eales, S. A., Taylor, G. B., & Bradley, A. D. 1998, MNRAS, 295, 265
 Blundell, K. M., Rawlings, S., & Willott, C. J. 1999, AJ, 117, 677
 Broderick, J. W., De Breuck, C., Hunstead, R. W., & Seymour, N. 2007, MNRAS, 375, 1059
 Cai, Z., Nan, R., Schilizzi, R. T., et al. 2002, A&A, 381, 401
 Carilli, C. L., Owen, F. N., & Harris, D. E. 1994, AJ, 107, 480
 Carilli, C. L., Roettgering, H. J. A., van Ojik, R., Miley, G. K., & van Breugel, W. J. M. 1997, ApJS, 109, 1
 Chambers, K. C., Miley, G. K., van Breugel, W. J. M., et al. 1996, ApJS, 106, 247
 Cimatti, A., di Serego-Alighieri, S., Fosbury, R. A. E., Salvati, M., & Taylor, D. 1993, MNRAS, 264, 421
 Cleary, K., Lawrence, C. R., Marshall, J. A., Hao, L., & Meier, D. 2007, ApJ, 660, 117
 Corbin, M. R., Charlot, S., De Young, D. S., Owen, F., & Dunlop, J. S. 1998, ApJ, 496, 803
 De Breuck, C., van Breugel, W., Röttgering, H. J. A., & Miley, G. 2000, A&AS, 143, 303
 De Breuck, C., van Breugel, W., Röttgering, H., et al. 2001, AJ, 121, 1241
 De Breuck, C., Seymour, N., Stern, D., et al. 2010, ApJ, 725, 36
 di Serego Alighieri, S., Fosbury, R. A. E., Tadhunter, C. N., & Quinn, P. J. 1989, Nature, 341, 307
 di Serego Alighieri, S., Cimatti, A., & Fosbury, R. A. E. 1993, ApJ, 404, 584
 Dullemond, C. P., & van Bemmelen, I. M. 2005, A&A, 436, 47
 Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, ApJS, 154, 10
 Ferrarese, L., & Merritt, D. 2000, ApJ, 539, L9
 Fioc, M., & Rocca-Volmerange, B. 1997, A&A, 326, 950
 Fitzpatrick, E. L. 1999, PASP, 111, 63
 Fomalont, E. B., Kellermann, K. I., Partridge, R. B., Windhorst, R. A., & Richards, E. A. 2002, AJ, 123, 2402
 Fritz, J., Franceschini, A., & Hatziminaoglou, E. 2006, MNRAS, 366, 767

- Giovannini, G., Feretti, L., Gregorini, L., & Parma, P. 1988, *A&A*, 199, 73
- Granato, G. L., & Danese, L. 1994, *MNRAS*, 268, 235
- Haas, M., Klaas, U., Müller, S. A. H., et al. 2003, *A&A*, 402, 87
- Haas, M., Willner, S. P., Heymann, F., et al. 2008, *ApJ*, 688, 122
- Hines, D. C. 1994, Ph.D. Thesis, Texas University
- Hoekstra, H., Barthel, P. D., & Hes, R. 1997, *A&A*, 319, 757
- Hönig, S. F., Beckert, T., Ohnaka, K., & Weigelt, G. 2006, *A&A*, 452, 459
- Houck, J. R., Roellig, T. L., van Cleve, J., et al. 2004, *ApJS*, 154, 18
- Humphrey, A., Villar-Martín, M., Vernet, J., et al. 2008, *MNRAS*, 383, 11
- Ibar, E., & Lira, P. 2007, *A&A*, 466, 531
- Kapahi, V. K. & Saikia, D. J. 1982, *J. Astrophys. Astron.*, 3, 465
- Kapahi, V. K., Athreya, R. M., van Breugel, W., McCarthy, P. J., & Subrahmanya, C. R. 1998, *ApJS*, 118, 275
- Kishimoto, M., Hönig, S. F., Antonucci, R., et al. 2011, *A&A*, 536, A78
- Krolik, J. H., & Begelman, M. C. 1988, *ApJ*, 329, 702
- Laing, R. A., Parma, P., de Ruiter, H. R., & Fantí, R. 1999, *MNRAS*, 306, 513
- Law-Green, J. D. B., Eales, S. A., Leahy, J. P., Rawlings, S., & Lacy, M. 1995a, *MNRAS*, 277, 995
- Law-Green, J. D. B., Leahy, J. P., Alexander, P., et al. 1995b, *MNRAS*, 274, 939
- Leipski, C., Haas, M., Willner, S. P., et al. 2010, *ApJ*, 717, 766
- Matthews, T. A., Morgan, W. W., & Schmidt, M. 1964, *ApJ*, 140, 35
- McCarthy, P. J., van Breugel, W., Spinrad, H., & Djorgovski, S. 1987, *ApJ*, 321, L29
- McCarthy, P. J., Spinrad, H., Dickinson, M., et al. 1990, *ApJ*, 365, 487
- McCarthy, P. J., van Breugel, W., & Kapahi, V. K. 1991, *ApJ*, 371, 478
- McLure, R. J., Jarvis, M. J., Targett, T. A., Dunlop, J. S., & Best, P. N. 2006, *MNRAS*, 368, 1395
- Mulchaey, J. S., Myshotzky, R. F., & Weaver, K. A. 1992, *ApJ*, 390, L69
- Mullin, L. M., & Hardcastle, M. J. 2009, *MNRAS*, 398, 1989
- Nenkova, M., Ivezić, Ž., & Elitzur, M. 2002, *ApJ*, 570, L9
- Nenkova, M., Sirocky, M. M., Ivezić, Ž., & Elitzur, M. 2008, *ApJ*, 685, 147
- Nesvadba, N. P. H., Lehnert, M. D., De Breuck, C., Gilbert, A. M., & van Breugel, W. 2008, *A&A*, 491, 407
- Nesvadba, N. P. H., De Breuck, C., Lehnert, M. D., et al. 2011, *A&A*, 525, A43
- Ogle, P. M., Cohen, M. H., Miller, J. S., et al. 1999, *ApJS*, 125, 1
- Pentericci, L., Van Reeven, W., Carilli, C. L., Röttgering, H. J. A., & Miley, G. K. 2000, *A&AS*, 145, 121
- Pentericci, L., McCarthy, P. J., Röttgering, H. J. A., et al. 2001, *ApJS*, 135, 63
- Pérez-Torres, M.-A., & De Breuck, C. 2005, *MNRAS*, 363, L41
- Pier, E. A., & Krolik, J. H. 1992, *ApJ*, 401, 99
- Pier, E. A., & Krolik, J. H. 1993, *ApJ*, 418, 673
- Ramos Almeida, C., Levenson, N. A., Rodríguez Espinosa, J. M., et al. 2009, *ApJ*, 702, 1127
- Rees, M. J. 1984, *ARA&A*, 22, 471
- Reuland, M., van Breugel, W., Röttgering, H., et al. 2003, *ApJ*, 592, 755
- Richards, G. T., Lacy, M., Storrie-Lombardi, L. J., et al. 2006, *ApJS*, 166, 470
- Rieke, G. H., Young, E. T., Engelbracht, C. W., et al. 2004, *ApJS*, 154, 25
- Rigby, E. E., Snellen, I. A. G., & Best, P. N. 2007, *MNRAS*, 380, 1449
- Rowan-Robinson, M. 1995, *MNRAS*, 272, 737
- Sanders, D. B., Soifer, B. T., Elias, J. H., et al. 1988, *ApJ*, 325, 74
- Schartmann, M., Meisenheimer, K., Camenzind, M., Wolf, S., & Henning, T. 2005, *A&A*, 437, 861
- Scheuer, P. A. G., & Readhead, A. C. S. 1979, *Nature*, 277, 182
- Seymour, N., Stern, D., De Breuck, C., et al. 2007, *ApJS*, 171, 353
- Seymour, N., Ogle, P., De Breuck, C., et al. 2008, *ApJ*, 681, L1
- Seymour, N., Altieri, B., De Breuck, C., et al. 2012, *ApJ*, 755, 146
- Smail, I., Swinbank, A. M., Ivison, R. J., & Ibar, E. 2011, *MNRAS*, 414, L95
- Spinrad, H., Marr, J., Aguilar, L., & Djorgovski, S. 1985, *PASP*, 97, 932
- Stalewski, M., Fritz, J., Baes, M., Nakos, T., & Popović, L. Č. 2012, *MNRAS*, 420, 2756
- Tran, H. D., Miller, J. S., & Kay, L. E. 1992, *ApJ*, 397, 452
- van Bemmell, I. M., & Dullemond, C. P. 2003, *A&A*, 404, 1
- van Breugel, W. J. M., Stanford, S. A., Spinrad, H., Stern, D., & Graham, J. R. 1998, *ApJ*, 502, 614
- van der Wolk, G., Barthel, P. D., Peletier, R. F., & Pel, J. W. 2010, *A&A*, 511, A64
- van Ojik, R., Roettgering, H. J. A., Carilli, C. L., et al. 1996, *A&A*, 313, 25
- Vernet, J., Fosbury, R. A. E., Villar-Martín, M., et al. 2001, *A&A*, 366, 7
- Villar-Martín, M., Vernet, J., di Serego Alighieri, S., et al. 2003, *MNRAS*, 346, 273
- Werner, M. W., Roellig, T. L., Low, F. J., et al. 2004, *ApJS*, 154, 1
- White, R. L., Becker, R. H., Helfand, D. J., & Gregg, M. D. 1997, *ApJ*, 475, 479
- Willott, C. J., Rawlings, S., Blundell, K. M., & Lacy, M. 2000, *MNRAS*, 316, 449

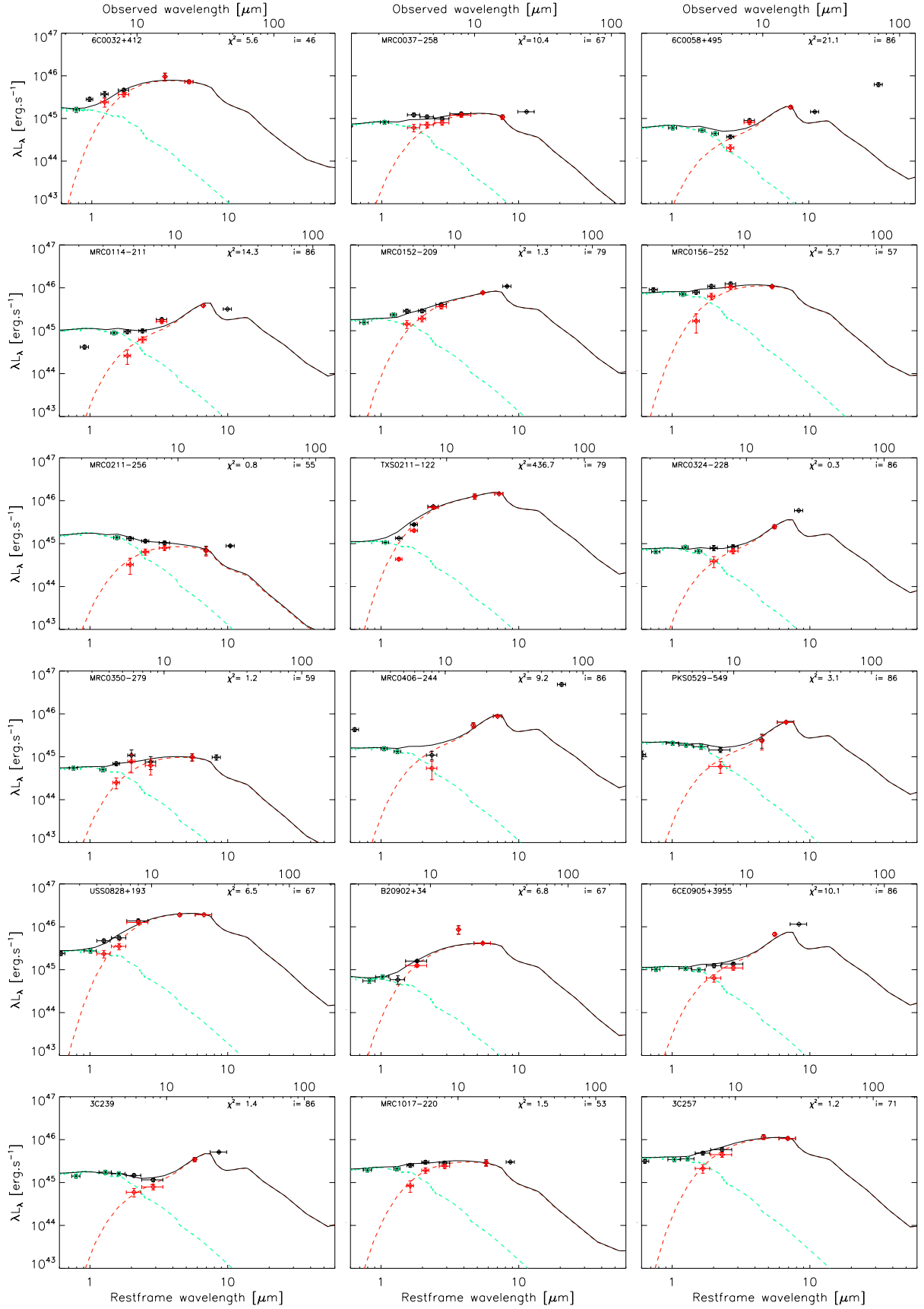


Fig. 11. Mid-IR SEDs for the sHzRG sample. Black points: photometric data. Green points: measurements used to normalise the contribution from an old stellar population. Red points: results after subtraction of the stellar contribution. Dashed green line: model stellar emission. Dashed red line: best-fitting torus model. Full black line: sum of stellar and torus models.

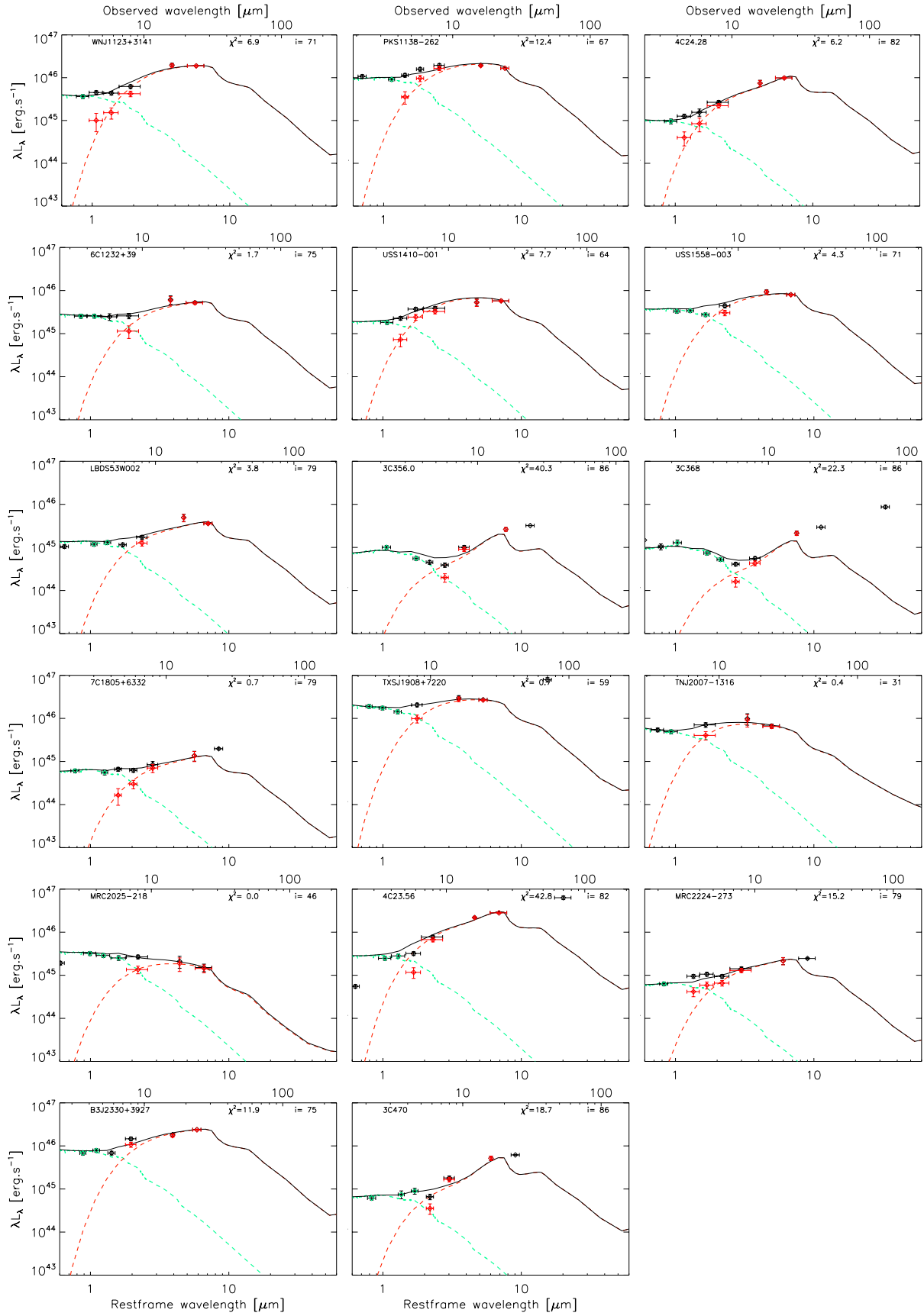


Fig. 11. continued.

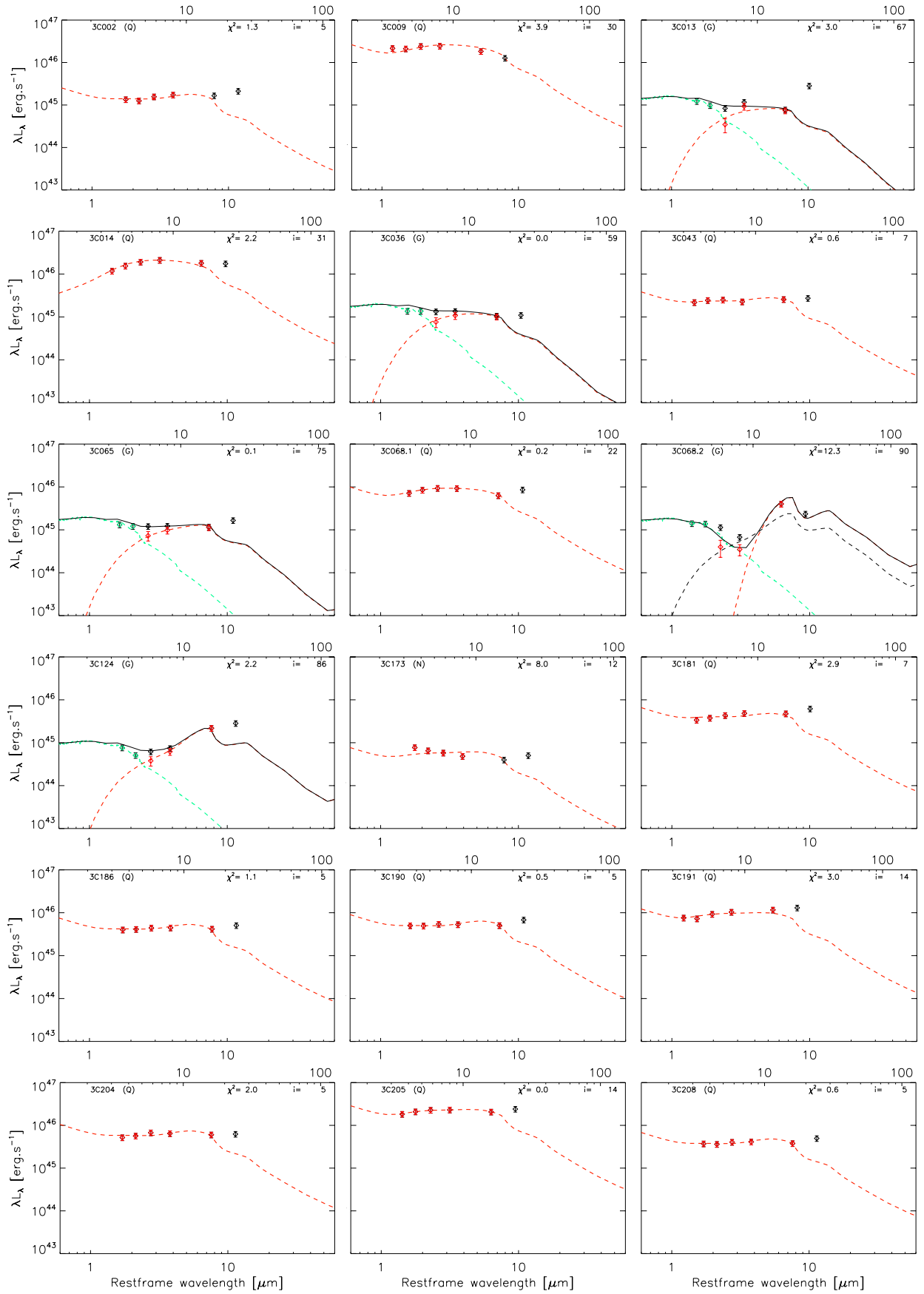


Fig. 12. Mid-IR SEDs for the s3CR sample. The symbols are the same than in Fig. 11.

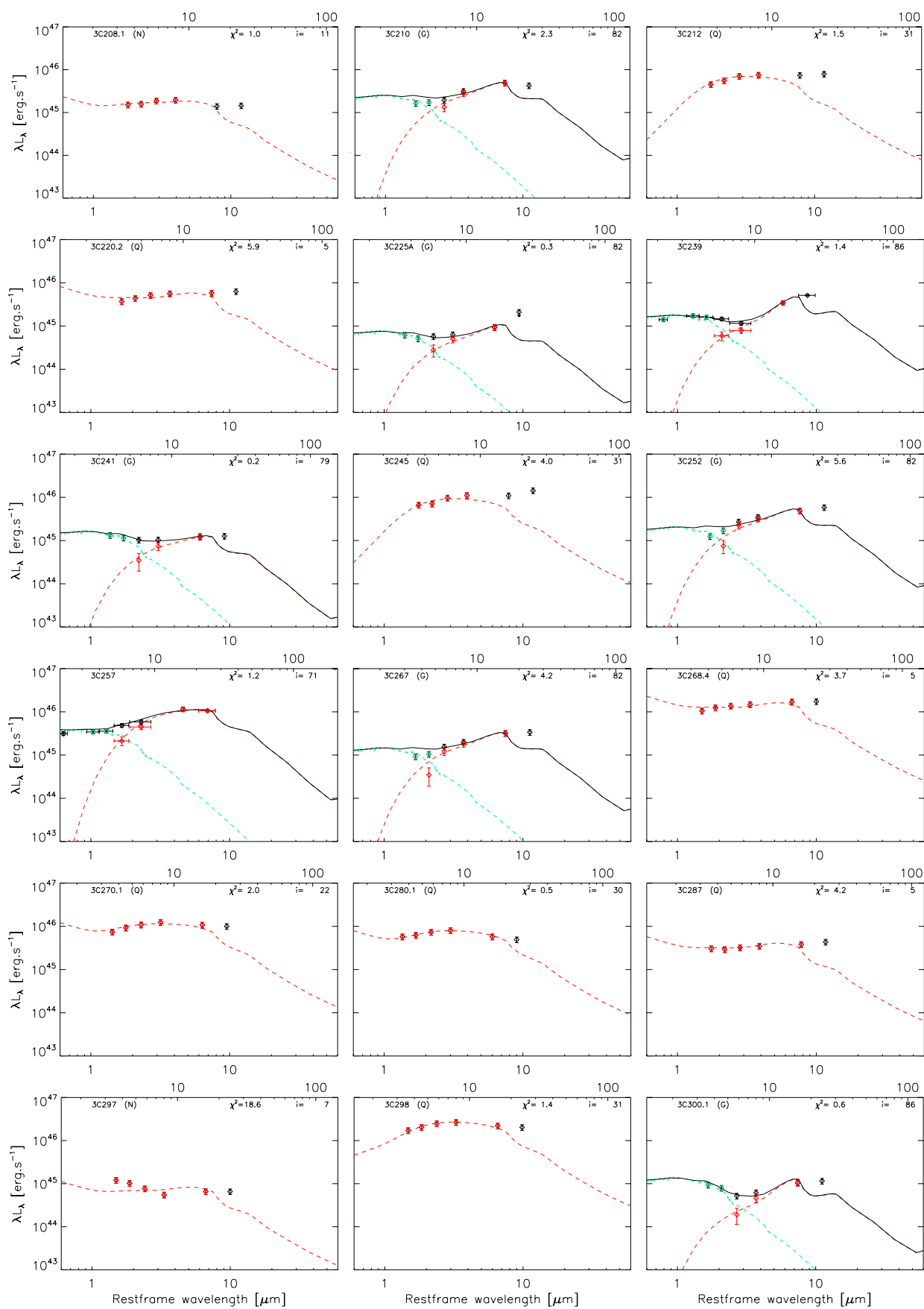


Fig. 12. continued.

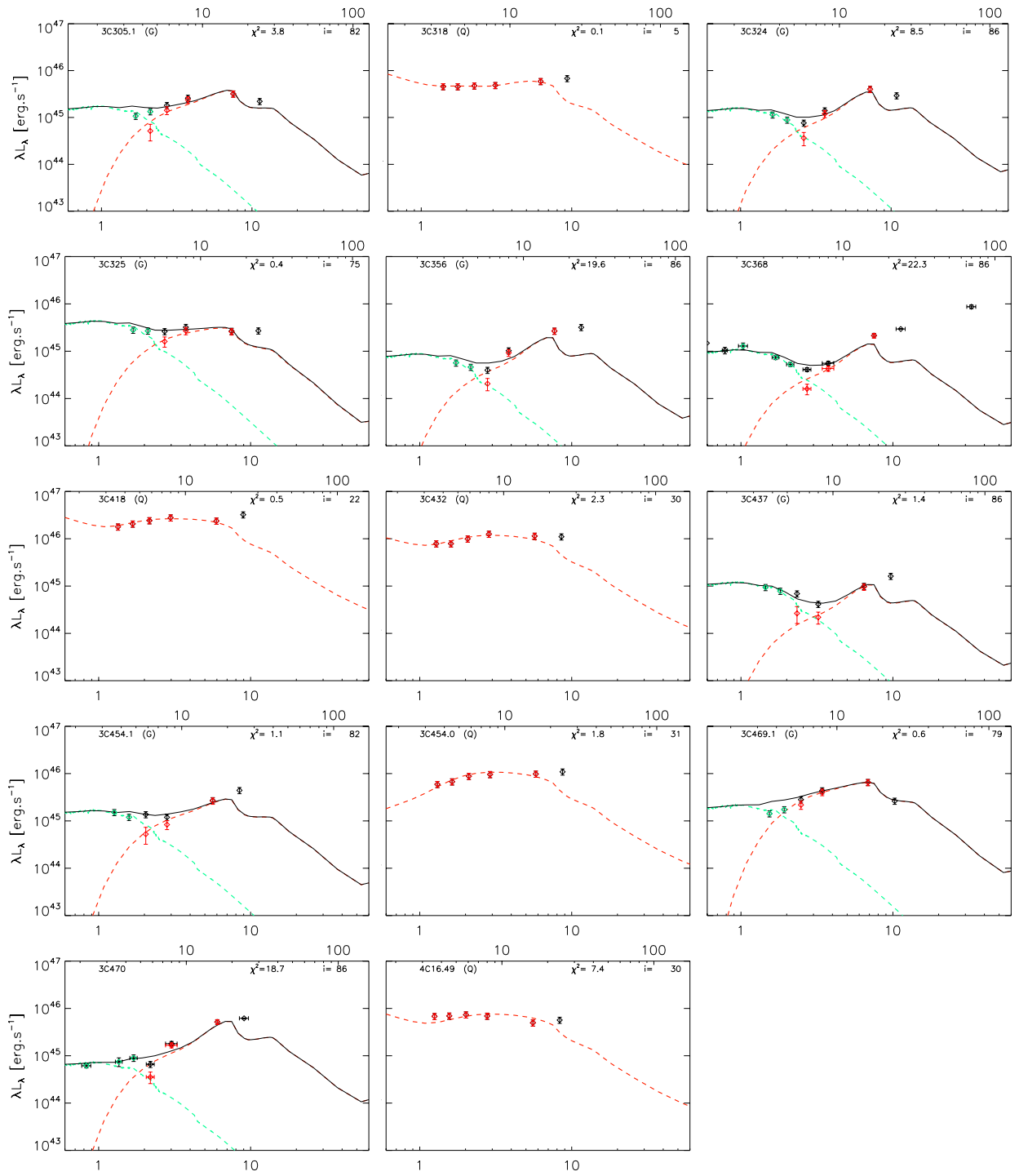


Fig. 12. continued.

Table 1. Radio data for the SHzRG sample from the literature and core dominance calculated.

Name	z	$\log P_{\text{extended}}^{500\text{ MHz}}$ [W Hz ⁻¹]	$S_{\text{core}}^{8.4\text{ GHz}}$ [mJy]	$\alpha_{4.8}^{8.4}$ core	R	Morph	References
6C0032+412	3.670	28.75	0.25	...	0.00020	T	Blundell et al. (1998)
MRC0037-258	1.100	27.72	1.74	0.9	0.0012	T	De Breuck et al. (2010); Kapahi et al. (1998)
6C0058+495	1.173	27.33	<0.00055	...	<0.00085*	D	Blundell et al. (1998)
MRC0114-211	1.410	28.66	1.60	...	0.00018	T	De Breuck et al. (2010); Kapahi et al. (1998)
TNJ0121+1320	3.516	28.49	D	De Breuck et al. (2000)
6C0132+330	1.710	27.64	1.63	0.1	0.0027	T	De Breuck et al. (2010)
6C0140+326	4.413	28.73	D	Blundell et al. (1998)
MRC0152-209	1.920	28.20	<0.00008	...	<5.3e-05*	D	Pentericci et al. (2000)
MRC0156-252	2.016	28.46	6.58	-0.7	0.0026	T	Carilli et al. (1997)
TNJ0205+2242	3.506	28.46	0.10	-0.6	0.00013*	T	De Breuck et al. (2000)
MRC0211-256	1.300	27.78	S	De Breuck et al. (2000)
TXS0211-122	2.340	28.48	1.23	-1.0	0.00073	T	Carilli et al. (1997)
3C65	1.176	28.63	0.52	-0.2	4.3e-05	T	Carilli et al. (1997); Corbin et al. (1998)
MRC0251-273	3.160	28.54	1.35	-1.6	0.0020	T	De Breuck et al. (2010); Kapahi et al. (1998)
MRC0316-257	3.130	28.95	0.28	-0.6	9.0e-05*	T	McCarthy et al. (1991); Athreya et al. (1997)
MRC0324-228	1.894	28.49	<0.00080	...	<0.00026	T	De Breuck et al. (2010); McCarthy et al. (1991)
MRC0350-279	1.900	28.25	0.19	...	0.00011*	T	De Breuck et al. (2010); Kapahi et al. (1998)
MRC0406-244	2.427	29.03	1.59	-0.8	0.00027	T	Carilli et al. (1997)
4C60.07	3.788	29.20	0.21	-1.7	0.00012	T	Carilli et al. (1997)
PKS0529-549	2.575	29.16	D	Broderick et al. (2007)
WNJ0617+5012	3.153	28.02	D	De Breuck et al. (2000)
4C41.17	3.792	29.18	0.27	-0.2	5.9e-05	T	Carilli et al. (1994)
WNJ0747+3654	2.992	28.14	S	De Breuck et al. (2000)
6CE0820+3642	1.860	28.28	0.29	-1.2	0.00016	T	De Breuck et al. (2010); Law-Green et al. (1995b)
5C7.269	2.218	27.82	D	De Breuck et al. (2010)
USS0828+193	2.572	28.44	2.81	-0.7	0.0020	T	Carilli et al. (1997)
6CE0901+3551	1.910	28.19	0.14	...	9.4e-05	T	De Breuck et al. (2010)
B20902+34	3.395	28.78	8.90	-0.1	0.0037	T	Carilli et al. (1994)
6CE0905+3955	1.883	28.17	0.33	...	0.00022	T	Law-Green et al. (1995a)
TNJ0924-2201	5.195	29.51	D	De Breuck et al. (2000)
6C0930+389	2.395	28.41	0.29	-0.8	0.00020	T	Pentericci et al. (2000)
USS0943-242	2.923	28.62	<0.00008	...	<5.3e-05*	D	Carilli et al. (1997)
3C 239	1.781	29.00	0.44	...	3.9e-05	T	Best et al. (1997)
MG1019+0534	2.765	28.57	2.02	-1.0	0.0015	T	Pentericci et al. (2000)
MRC1017-220	1.768	27.94	S	Pentericci et al. (2000)
WNJ1115+5016	2.540	27.82	D	De Breuck et al. (2000)
3C 257	2.474	29.16	<0.00090	...	<0.00012*	D	van Breugel et al. (1998)
WNJ1123+3141	3.217	28.51	0.93	...	0.00098	T	White et al. (1997)
PKS1138-262	2.156	29.07	1.98	-1.3	0.00027	T	Carilli et al. (1997)
3C 266	1.275	28.54	<0.00018	...	<2.1e-05	D	Best et al. (1997)
6C1232+39	3.220	28.93	0.35	-0.2	0.00010	T	Carilli et al. (1997)
USS1243+036	3.570	29.23	0.70	-1.0	0.00020	T	van Ojik et al. (1996)
TNJ1338-1942	4.110	28.71	0.16	-1.0	0.00021	T	Pentericci et al. (2000)
4C24.28	2.879	29.05	0.51	-0.6	0.00011	T	Carilli et al. (1997)
3C 294	1.786	28.96	T	McCarthy et al. (1990)
USS1410-001	2.363	28.41	2.27	-0.7	0.0015	T	Carilli et al. (1997)
8C1435+635	4.250	29.40	1.31	-1.3	0.00049	T	Carilli et al. (1997)
USS1558-003	2.527	28.82	0.97	-0.1	0.00022	T	Pentericci et al. (2000)
USS1707+105	2.349	28.63	D	Pentericci et al. (2001)
LBDS53W002	2.393	27.78	S	Fomalont et al. (2002)
LBDS53W091	1.552	27.04	D	Rigby et al. (2007)
3C 356	1.079	28.35	0.22	...	2.7e-05	T	Best et al. (1997)
7C1751+6809	1.540	27.46	D	De Breuck et al. (2000)
7C1756+6520	1.416	27.40	D	Rigby et al. (2007)
3C 368	1.132	28.52	0.16	-0.5	1.5e-05	T	Best et al. (1997)
7C1805+6332	1.840	27.78	0.15	-1.2	0.00025	T	De Breuck et al. (2010)
4C40.36	2.265	28.79	<0.00008	...	<2.0e-05	D	Carilli et al. (1997)
TXSJ1908+7220	3.530	29.12	4.80	-1.3	0.0021	T	Pentericci et al. (2000)
WNJ1911+6342	3.590	28.14	S	De Breuck et al. (2000)
TNJ2007-1316	3.840	29.13	1.60	...	0.00060	T	De Breuck et al. (2010)
MRC2025-218	2.630	28.74	0.38	-0.9	0.00016	T	Carilli et al. (1997)
MRC2048-272	2.060	28.72	<0.00008	...	<1.9e-05	D	Pentericci et al. (2000)
MRC2104-242	2.491	28.84	0.19	-1.6	7.2e-05	T	Pentericci et al. (2000)
4C23.56	2.483	28.93	4.95	-0.9	0.0012	T	Carilli et al. (1997)
MG2144+1928	3.592	29.08	<0.00008	...	<2.9e-05*	D	Carilli et al. (1997)

Table 1. continued.

Name	z	$\log P_{\text{extended}}^{500\text{MHz}}$ [W Hz ⁻¹]	$S_{\text{core}}^{8.4\text{GHz}}$ [mJy]	$\alpha_{4.8}^{8.4}$ core	R	Morph	References
USS2202+128	2.706	28.54	0.14	-2.5	0.00020	T	Carilli et al. (1997)
MRC2224-273	1.679	27.52	S	Pentericci et al. (2000)
B3J2330+3927	3.086	28.33	3.98	-0.1	0.0040	T	Pérez-Torres & De Breuck (2005)
4C28.58	2.891	28.91	0.70	-0.2	0.00017	T	Cai et al. (2002) ; Chambers et al. (1996)
3C 470	1.653	28.79	1.20	...	0.00014	T	Best et al. (1997)

Notes. ^(*) Mark the core flux recalculated in this paper, see Sect. 2.2. The letters S, D and T refer to the morphology of the radio source, with 1, 2 or 3 identified components, respectively. Names in bold are the sSHzRG sample.

Table 2. 3CR sample, *Spitzer* observations.

Name	Type	z	3.6 μm [μJy]	4.5 μm [μJy]	5.8 μm [μJy]	8 μm [μJy]	16 μm [μJy]	24 μm [μJy]
3C 002	Q	1.04	283 \pm 42	330 \pm 50	530 \pm 80	809 \pm 121	1550 \pm 233	2970 \pm 446
3C 009	Q	2.01	884 \pm 133	1080 \pm 162	1590 \pm 239	2220 \pm 333	3330 \pm 500	3470 \pm 520
3C 013	G	1.35	133 \pm 20	133 \pm 20	147 \pm 22	283 \pm 42	375 \pm 56	2060 \pm 309
3C 014	Q	1.47	1040 \pm 156	1710 \pm 257	2740 \pm 411	4150 \pm 623	7070 \pm 1061	10300 \pm 1545
3C 036	G	1.30	163 \pm 24	205 \pm 31	256 \pm 38	360 \pm 54	560 \pm 84	874 \pm 131
3C 043	Q	1.47	193 \pm 29	270 \pm 41	356 \pm 53	445 \pm 67	1010 \pm 152	1610 \pm 242
3C 065	G	1.18	202 \pm 30	233 \pm 35	299 \pm 45	418 \pm 63	798 \pm 120	1700 \pm 255
3C 068.1	Q	1.24	967 \pm 145	1430 \pm 215	2040 \pm 306	2780 \pm 417	3800 \pm 570	7760 \pm 1164
3C 068.2	G	1.58	105 \pm 16	129 \pm 19	137 \pm 21	112 \pm 17	1340 \pm 201	1170 \pm 176
3C 119	G	1.02	802 \pm 120	878 \pm 132	1280 \pm 192	1850 \pm 278	4820 \pm 723	8260 \pm 1239
3C 124	G	1.08	144 \pm 22	120 \pm 18	188 \pm 28	310 \pm 47	1840 \pm 276	3560 \pm 534
3C 173	N	1.03	163 \pm 24	172 \pm 26	197 \pm 30	227 \pm 34	374 \pm 56	710 \pm 107
3C 181	Q	1.38	348 \pm 52	485 \pm 73	722 \pm 108	1110 \pm 167	2180 \pm 327	4260 \pm 639
3C 186	Q	1.06	791 \pm 119	1020 \pm 153	1410 \pm 212	1960 \pm 294	3660 \pm 549	6660 \pm 999
3C 190	Q	1.20	739 \pm 111	908 \pm 136	1290 \pm 194	1740 \pm 261	3310 \pm 497	6690 \pm 1004
3C 191	Q	1.96	333 \pm 50	399 \pm 60	655 \pm 98	1010 \pm 152	2270 \pm 341	3810 \pm 572
3C 194	G	1.78	201 \pm 30	176 \pm 26	164 \pm 25	208 \pm 31	509 \pm 76	885 \pm 133
3C 204	Q	1.11	917 \pm 138	1250 \pm 188	1920 \pm 288	2540 \pm 381	4730 \pm 710	7360 \pm 1104
3C 205	Q	1.53	1460 \pm 219	2080 \pm 312	2920 \pm 438	4090 \pm 614	7320 \pm 1098	12800 \pm 1920
3C 208	Q	1.11	660 \pm 99	803 \pm 120	1160 \pm 174	1620 \pm 243	2980 \pm 447	5870 \pm 881
3C 208.1	N	1.02	331 \pm 50	430 \pm 65	656 \pm 98	954 \pm 143	1360 \pm 204	2110 \pm 317
3C 210	G	1.17	256 \pm 38	336 \pm 50	489 \pm 73	1090 \pm 164	3410 \pm 512	4430 \pm 665
3C 212	Q	1.05	925 \pm 139	1430 \pm 215	2340 \pm 351	3400 \pm 510	6710 \pm 1007	10800 \pm 1620
3C 220.2	Q	1.16	592 \pm 89	870 \pm 131	1330 \pm 200	2000 \pm 300	4150 \pm 623	6720 \pm 1008
3C 222	G	1.34	83 \pm 12	91 \pm 14	73 \pm 11	65 \pm 10	331 \pm 50	229 \pm 34
3C 225A	G	1.56	47 \pm 7	49 \pm 7	71 \pm 11	108 \pm 16	321 \pm 48	<1070
3C 230	G	1.49	1040 \pm 156	672 \pm 101	438 \pm 66	317 \pm 48	1150 \pm 173	1560 \pm 234
3C 238	G	1.40	65 \pm 10	77 \pm 12	84 \pm 12	<92	<283	266 \pm 40
3C 239	G	1.78	96 \pm 14	111 \pm 17	130 \pm 20	142 \pm 21	651 \pm 98	1450 \pm 218
3C 241	G	1.62	92 \pm 14	101 \pm 15	116 \pm 17	161 \pm 24	389 \pm 58	591 \pm 89
3C 245	Q	1.03	1420 \pm 213	1900 \pm 285	3350 \pm 503	5270 \pm 790	10400 \pm 1560	20400 \pm 3060
3C 249	G	1.55	54 \pm 8	52 \pm 8	42 \pm 6	47 \pm 7	194 \pm 29	<516
3C 250	G	1.26	61 \pm 9	59 \pm 9	46 \pm 7	29 \pm 4	162 \pm 24	<147
3C 252	G	1.10	225 \pm 34	382 \pm 57	787 \pm 118	1390 \pm 209	3900 \pm 585	7000 \pm 1050
3C 255	G	1.36	85 \pm 13	86 \pm 13	57 \pm 9	22 \pm 3	<116	<241
3C 256	G	1.82	34 \pm 5	37 \pm 6	43 \pm 7	75 \pm 11	743 \pm 111	1900 \pm 285
3C 257	G	2.47	85 \pm 13	111 \pm 17	194 \pm 29	322 \pm 48	...	1360 \pm 204
3C 266	G	1.27	68 \pm 10	73 \pm 11	45 \pm 7	102 \pm 15	370 \pm 56	980 \pm 147
3C 267	G	1.14	153 \pm 23	218 \pm 33	414 \pm 62	739 \pm 111	2370 \pm 356	3730 \pm 560
3C 268.4	Q	1.40	1060 \pm 159	1560 \pm 234	2220 \pm 333	3330 \pm 500	7580 \pm 1137	11600 \pm 1740
3C 270.1	Q	1.52	606 \pm 91	944 \pm 142	1430 \pm 214	2260 \pm 339	3910 \pm 587	5470 \pm 821
3C 280.1	Q	1.66	378 \pm 57	512 \pm 77	777 \pm 116	1170 \pm 176	1680 \pm 252	2160 \pm 324
3C 287	Q	1.05	613 \pm 92	735 \pm 110	1050 \pm 157	1560 \pm 234	3430 \pm 515	5820 \pm 873
3C 294	G	1.79	<93	<103	68 \pm 10	67 \pm 10	...	348 \pm 52
3C 297	N	1.41	119 \pm 18	126 \pm 19	122 \pm 18	121 \pm 18	<288	432 \pm 65
3C 298	Q	1.44	1600 \pm 240	2390 \pm 359	3710 \pm 556	5510 \pm 827	9160 \pm 1374	12600 \pm 1890
3C300.1	G	1.16	148 \pm 22	158 \pm 24	133 \pm 20	220 \pm 33	751 \pm 113	1220 \pm 183
3C 305.1	G	1.13	181 \pm 27	282 \pm 42	495 \pm 74	972 \pm 146	2410 \pm 362	2490 \pm 374
3C 318	Q	1.57	343 \pm 51	427 \pm 64	571 \pm 86	806 \pm 121	1960 \pm 294	3400 \pm 510
3C 322	G	1.68	128 \pm 19	135 \pm 20	94 \pm 14	120 \pm 18	411 \pm 62	804 \pm 121
3C 324	G	1.21	165 \pm 25	160 \pm 24	178 \pm 27	450 \pm 68	2580 \pm 387	2820 \pm 423
3C 325	G	1.13	472 \pm 71	565 \pm 85	708 \pm 106	1200 \pm 180	1990 \pm 299	3030 \pm 455
3C 326.1	G	1.83	29 \pm 4	34 \pm 5	26 \pm 4	72 \pm 11	829 \pm 124	1430 \pm 215
3C 356	G	1.08	108 \pm 16	110 \pm 16	122 \pm 18	434 \pm 65	2270 \pm 341	4060 \pm 609
3C 368	G	1.13	126 \pm 19	112 \pm 17	112 \pm 17	210 \pm 32	1370 \pm 206	3250 \pm 488
3C 418	Q	1.69	1130 \pm 170	1630 \pm 245	2470 \pm 371	3900 \pm 585	6680 \pm 1002	13600 \pm 2040
3C 432	Q	1.80	420 \pm 63	526 \pm 79	857 \pm 129	1490 \pm 224	2710 \pm 407	3940 \pm 591
3C 437	G	1.48	82 \pm 12	85 \pm 13	97 \pm 15	80 \pm 12	384 \pm 58	941 \pm 141
3C 454.1	G	1.84	77 \pm 12	76 \pm 11	112 \pm 17	135 \pm 20	612 \pm 92	1500 \pm 225
3C 454.0	Q	1.76	339 \pm 51	481 \pm 72	811 \pm 122	1220 \pm 183	2490 \pm 374	4150 \pm 623
3C 469.1	G	1.34	160 \pm 24	244 \pm 37	509 \pm 76	1090 \pm 164	3270 \pm 491	1970 \pm 296
3C 470	G	1.65	50 \pm 7	75 \pm 11	72 \pm 11	266 \pm 40	1510 \pm 227	2650 \pm 398
4C 13.66	G	1.45	24 \pm 4	24 \pm 4	21 \pm 3	18 \pm 3	<260	276 \pm 41
4C 16.49	Q	1.88	329 \pm 49	420 \pm 63	573 \pm 86	743 \pm 111	1070 \pm 161	1830 \pm 275

Notes. Q = quasars, G = radio galaxies, N = undefined, Names in bold correspond to the s3CR sample.

Table 6. Results for the s3CR sample, from Sects. 4.3 and 4.4.

Name	Type	A_v	i [$^\circ$]	$\log \nu P_{\nu(5\ \mu\text{m})}^{\text{AGN}}$ [W]
3C 002	Q	-1.5	5	38.25
3C 009	Q	-3.1	30	39.36
3C 013	G	6.2	67	37.92
3C 014	Q	-1.5	31	39.27
3C 036	G	1.5	59	38.07
3C 043	Q	-3.1	7	38.45
3C 065	G	6.2	75	38.08
3C 068.1	Q	-4.7	22	38.93
3C 068.2	G	68.2	89	38.35
3C 124	G	38.8	86	38.09
3C 173	N	-10.8	12	37.79
3C 181	Q	-1.5	7	38.68
3C 186	Q	-3.1	5	38.73
3C 190	Q	-4.7	5	38.81
3C 191	Q	-1.5	14	39.00
3C 204	Q	-3.1	5	38.87
3C 205	Q	-3.1	14	39.37
3C 208	Q	-3.1	5	38.68
3C 208.1	N	-1.5	11	38.26
3C 210	G	21.7	82	38.56
3C 212	Q	0.0	31	38.79
3C 220.2	Q	0.0	5	38.76
3C 225A	G	15.5	82	37.89
3C 239	G	27.9	86	38.33
3C 241	G	13.9	79	38.03
3C 245	Q	1.5	31	38.92
3C 252	G	18.6	82	38.59
3C 257	G	6.2	67	38.96
3C 267	G	20.1	82	38.38
3C 268.4	Q	-1.5	5	39.21
3C 270.1	Q	-1.5	22	39.02
3C 280.1	Q	-3.1	30	38.84
3C 287	Q	-1.5	5	38.61
3C 297	N	-10.8	7	37.92
3C 298	Q	-1.5	31	39.37
3C 300.1	G	29.4	86	37.86
3C 305.1	G	17.0	82	38.44
3C 318	Q	-3.1	5	38.77
3C 324	G	41.8	86	38.30
3C 325	G	6.2	75	38.46
3C 356	G	48.0	86	38.05
3C 368	G	31.0	86	38.01
3C 418	Q	-1.5	22	39.39
3C 432	Q	-1.5	30	39.01
3C 437	G	38.8	86	37.79
3C 454.1	G	23.2	82	38.32
3C 454.0	Q	-1.5	31	38.97
3C 469.1	G	15.5	79	38.75
3C 470	G	37.2	86	38.39
4C 16.49	Q	-4.7	30	38.82

Notes. Results for the s3CR sample. Types: Q = quasar, G = radio galaxy and N = unidentified.