



HAL
open science

A major electronics upgrade for the H.E.S.S. Cherenkov telescopes 1-4

G. Giavitto, T. Ashton, A. Balzer, D. Berge, F. Brun, T. Chaminade, E. Delagnes, G. Fontaine, M. Füssling, B. Giebels, et al.

► **To cite this version:**

G. Giavitto, T. Ashton, A. Balzer, D. Berge, F. Brun, et al.. A major electronics upgrade for the H.E.S.S. Cherenkov telescopes 1-4. 34th International Cosmic Ray Conference, Jul 2015, The Hague, Netherlands. pp.996, 10.22323/1.236.0996 . hal-03627388

HAL Id: hal-03627388

<https://hal.science/hal-03627388v1>

Submitted on 6 Sep 2024

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.



Distributed under a Creative Commons Attribution - NonCommercial - ShareAlike 4.0 International License

A major electronics upgrade for the H.E.S.S. Cherenkov telescopes 1-4

G. Giavitto^{*a}, T. Ashton^b, A. Balzer^c, D. Berge^c, F. Brun^d, T. Chaminade^d, E. Delagnes^d, G. Fontaine^f, M. Füßling^a, B. Giebels^f, J.F. Glicenstein^d, T. Gräber^a, J.A. Hinton^{b,g}, A. Jahnke^g, S. Klepser^a, M. Kossatz^a, A. Kretzschmann^a, V. Lefranc^{a,e}, H. Leich^a, H. Lüdecke^a, P. Manigot^f, V. Marandon^g, E. Moulin^d, M. de Naurois^f, P. Nayman^e, M. Penno^a, D. Ross^b, D. Salek^c, M. Schade^a, T. Schwab^g, R. Simoni^c, C. Stegmann^a, J. Thornhill^b, F. Toussenel^e

^aDESY, D-15738 Zeuthen, Germany

^bDepartment of Physics and Astronomy, The University of Leicester, University Road, Leicester, LE1 7RH, United Kingdom

^cGRAPPA, Anton Pannekoek Institute for Astronomy, University of Amsterdam, Science Park 904, 1098 XH Amsterdam, The Netherlands

^dDSM/Irfu, CEA Saclay, F-91191 Gif-Sur-Yvette Cedex, France

^eLPNHE, Université Pierre et Marie Curie Paris 6, Université Denis Diderot Paris 7, CNRS/IN2P3, 4 Place Jussieu, F-75252, Paris Cedex 5, France

^fLaboratoire Leprince-Ringuet, Ecole Polytechnique, CNRS/IN2P3, F-91128 Palaiseau, France

^gMax-Planck-Institut für Kernphysik, P.O. Box 103980, D 69029 Heidelberg, Germany

Email: gianluca.giavitto@desy.de, stefan.klepser@desy.de

The High Energy Stereoscopic System (H.E.S.S.) is an array of imaging atmospheric Cherenkov telescopes (IACTs) located in the Khomas Highland in Namibia. It consists of four 12-m telescopes (CT1-4), which started operations in 2003, and a 28-m diameter one (CT5), which was brought online in 2012. It is the only IACT system featuring telescopes of different sizes, which provides sensitivity for gamma rays across a very wide energy range, from ~ 30 GeV up to ~ 100 TeV. Since the camera electronics of CT1-4 are much older than the one of CT5, an upgrade is being carried out; first deployment was in 2015, full operation is planned for 2016. The goals of this upgrade are threefold: reducing the dead time of the cameras, improving the overall performance of the array and reducing the system failure rate related to aging. Upon completion, the upgrade will assure the continuous operation of H.E.S.S. at its full sensitivity until and possibly beyond the advent of CTA. In the design of the new components, several CTA concepts and technologies were used and are thus being evaluated in the field: The upgraded read-out electronics is based on the NECTAR readout chips; the new camera front- and back-end control subsystems are based on an FPGA and an embedded ARM computer; the communication between subsystems is based on standard Ethernet technologies. These hardware solutions offer good performance, robustness and flexibility. The design of the new cameras is reported here.

The 34th International Cosmic Ray Conference,

30 July- 6 August, 2015

The Hague, The Netherlands

1. Motivation of the H.E.S.S. I Upgrade

The H.E.S.S. Cherenkov telescopes 1-4 (aka CT1-4, or "H.E.S.S. I array") were installed in Namibia, near the Gamsberg mountain, between 2002 and 2004 [1]. In 2012, a fifth telescope (CT5) was inaugurated in the middle of the previous H.E.S.S. I array. In contrast to CT1-4, this new 28-m telescope was not only equipped with more modern camera electronics, but also operates at a much lower threshold of ~ 30 GeV, and therefore delivers a much higher rate of air shower triggers. Therefore, as MAGIC [6] and VERITAS [7] in the past years, also H.E.S.S. is undergoing a substantial hardware upgrade to optimise its performance in its final years.

The main reasons to upgrade the old CT1-4 cameras [2] are (i) to reduce the downtime of individual cameras, induced by their ageing and increasing failure rates of cable connectors and other critical parts, and (ii) to reduce the dead time of the array when operated in coincidence mode with CT5. The readout time of the original CT1-4 cameras amounts to about $\sim 450 \mu\text{s}$, which was acceptable for the typical array trigger rates of H.E.S.S. I (200 – 300 Hz). However, if these telescopes trigger in coincidence with CT5, array trigger rates of 1.5 kHz or more are usual, which can lead to a substantial fraction of events that are not recorded in CT1-4, which effectively corresponds to a dead time for *stereoscopic* events that can amount to 30% or more.

The way to improve on this situation is to replace the current readout scheme with a fast, Ethernet-based readout and modern front-end boards built around the NECTAR analog readout chips [3]. To reduce the failure rate, novel schemes for cabling, ventilation, power supply and pneumatics were developed. The first upgraded camera is deployed in 2015.

These developments offer the opportunity to improve the performance of the array. For instance, since the new readout can sustain much higher trigger rates, it would be possible to reduce the trigger thresholds of the CT1-4, thereby extending their sensitivity to gamma-ray showers of energies below 100 GeV. Other characteristics of the upgraded trigger and readout could also improve the performance of CT1-4 at higher energies.

2. Components and aspects

The upgraded components include electronic and mechanical parts. Figure 1 summarizes schematically their architecture and interplay. Basically everything inside the camera has been replaced or refurbished, except the photomultiplier tubes (PMTs) and their bases, that generate the high voltage.

Following the signal path, from right to left in the figure, the first upgraded components are the analog front-end boards inside the 60 *drawer* modular units. Each drawer consists of 16 PMT pixels, whose signals are read out by two analog boards and is controlled by one slow control board. The back-end electronics, which was previously composed of many rack units in two electronics racks, is now deployed in one rack only. Most of the control, security and camera trigger logic is implemented on one single rack unit called *Drawer Interface Box* (DIB). The drawers are connected to it, to the data and power networks via a connection board hosting three connectors: two RJ45 connectors, one for ethernet and one for trigger, readout control and timing, and one M8 4-pin connector for power (delivered at 24 V DC); The connection board also hosts the main DC-DC

*Speaker.

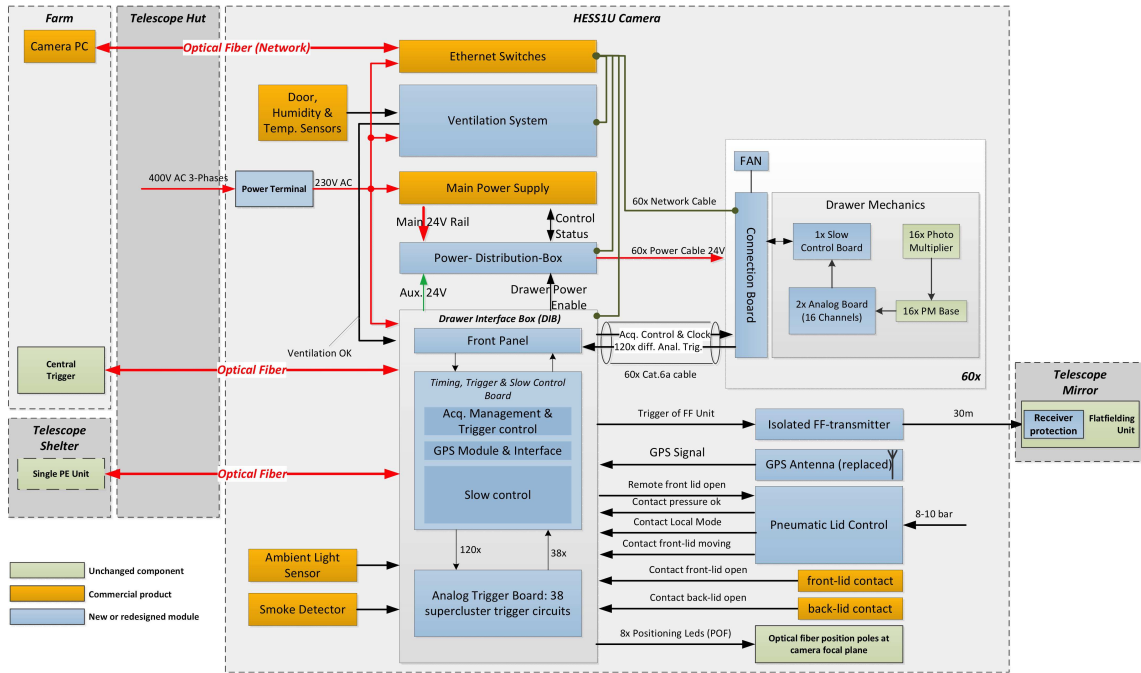


Figure 1: Electronics architecture of the H.E.S.S. I Upgrade. The original H.E.S.S. I systems are denoted by green boxes. Where possible, commercial components were used (in orange), also in the cabling, where either standard Ethernet or M8 cables were deployed. The dashed boxes correspond to physical locations: the bigger “HESS1U Camera“ box corresponds to the camera itself, the “telescope shelter” is the hut in which the camera is parked during the day, the “telescope hut” is mounted on the structure of the telescope (and co-rotates), the “farm” is located inside the control building.

converter of the drawer, in order to attain a galvanic separation between front- and back-end. The DIB is composed of 3 boards: a front panel, a main board and an analog trigger board. The front connection board is equipped with 60 RJ45 connectors for the drawers clock, acquisition control and trigger signals, an ethernet connector, several M8 connectors to sensors and actuators and auxiliary LEMO-00 connectors. The control and security functions are implemented on the FPGA on the main board. All peripheral components, like GPS, pneumatics, power supply, ventilation, are controlled from there, and all the sensors are read out from there. A security interlock logic on the FPGA firmware evaluates the sensor inputs: in case of errors or failures it shuts down the power of the drawers and closes the camera lid. The main board hosts a separated analog trigger board, responsible of the camera trigger decision. The drawers are powered by the *Power Distribution Box* (PDB), a 64 channel power switch which monitors constantly the current consumption of each drawer. The main power supply is a commercial 3-module unit which delivers the 24 V and ~ 80 A needed by the camera. Three modules share the load, but only two are actually needed, one serves as a hot spare. The DIB, the power switch and drawers are all equipped with a combination of a Cyclone IV FPGA and an ARM-based μ Computer Module (Stamp9G45), which allows for maximum flexibility and ease of control and communication.

The ventilation concept has been changed completely: a new ventilation system was imple-

mented on the new back door of the camera, with a single fan, replacing the many fans previously mounted on the camera sides. Pneumatic elements were added to be able to open and close the heavier back door, and the control unit of the entire camera pneumatics was renewed as well.

Many sensors were deployed inside and outside the camera for ambient light, temperature, smoke and humidity. Some of them are included in a new interlock system to guarantee safe camera operations for both shift crew and hardware.

2.1 Front-end electronics

The signal from the PMT base reaches the MCX connector on the analog board via a 15 cm long coaxial cable, and is there terminated with $R = 50\Omega$. The AC coupled signal is then pre-amplified, split into three branches and further amplified. Two of these three signal branches are the analog high and low gain, with total amplification factors 15.1 and 0.68, respectively. They are routed to the inputs of the NECTAR chip, which have a range of 2V. An adjustable offset is added to the signals at this point, in order to account for a possible undershoot and drifts in the baseline due to temperature effects. Its default value is about 205 mV. The third signal branch is amplified by a factor 45 and routed to a comparator, the output of which goes directly into the FPGA, where is sampled at 800 MHz and used to generate the trigger.

The NECTAR chip is a fast (1 GS/s) sampler and digitizer chip. It integrates three functions, namely (i) an analogue memory in which the analogue signal is sampled and stored at high rate, (ii) a 12-bit 21 MHz ADC digitizing the data stored in the analogue memory and (iii) a serializer, which sends the digital data downstream to the FPGA. Each chip integrates two channels of differential analogue memories with a depth of 1024 cells each. Operations are simultaneous on the 2 channels. In normal operations only a small fraction of the cells inside a region of interest (usually 16 cells long) are read out. The sampling is stopped by an external trigger signal, which must be synchronous to the analogue signal one wishes to digitize. When reading 16 cells, the nominal dead time of a NECTAR chip is lower than $\sim 2\mu\text{s}$. The way the chip is read out by the FPGA however requires a minimum safe interval between two events of $\sim 5.5\mu\text{s}$, so one can take that value as a measure of the overall dead time of the camera.

The typical bandwidth of the NECTAR chip analog inputs as reported in its datasheet [4] is between 300 and 500 MHz. Care was taken to use analog components matching that value. The overall end-to-end bandwidth of the readout was measured to be $\sim 330\text{MHz}$.

The analog part of the readout was developed to have very low noise. The pedestal noise, measured as the RMS of the value of a single cell is around 4 ADC counts, or 2 mV. At the nominal PMT gain of 2×10^5 , this corresponds to 0.32 photoelectrons.

The linearity of the high gain channel is better than 2%, whereas the low gain presents correlated deviations from linearity up to 8%; these errors can be corrected a posteriori since an absolute calibration of the transfer functions is performed prior to the installation on the telescope.

The cross talk between two channels within one NECTAR chip and between channels on different nectar chips is typically less than 0.5%, and never larger than 1%.

2.2 Trigger chain

The trigger scheme chosen for the upgraded camera is the same as for the current cameras [8]: An N -majority over overlapping sectors of 64 pixels each. As it has been already mentioned,

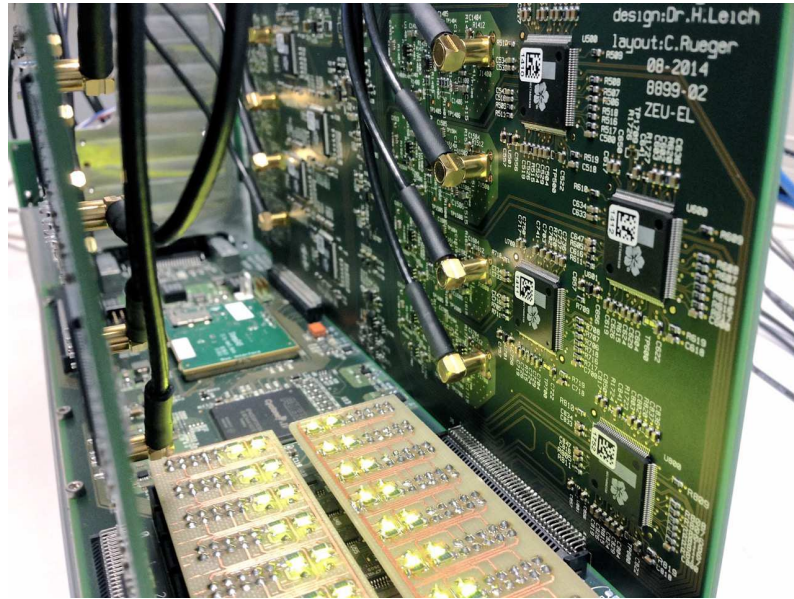


Figure 2: View into a new drawer, with analog front-end boards at the right and left, and drawer slow control board at the bottom. The analog boards contain the NECTAR chips (with the flower logo), the slow control board has test loads mounted at the plugs for the PMT bases (green lights).

the pixel comparator output is sampled by the FPGA on the drawer, which in turn generates two analog signals and sends them to the drawer interface box. These signals carry the number of pixels above the threshold on its left and right half, respectively; the number is passed as an analog pulse coded in height in steps of 33 mV. Once received by the drawer interface box, the signals are isochronously routed to the 38 overlapping trigger sectors on the analog trigger board, where an analog sum is built. The 60 drawers are arranged to occupy the central positions of a 9x8 matrix, and the sectors overlap in width by one half-drawer and in height by one full drawer, so the signal from each half drawer is routed to at most 4 sectors. The output of the each sector sum goes to a comparator, where a common threshold signal (corresponding to N pixels fired) is applied. The 38 comparator outputs are finally combined in an OR logic to give the camera trigger.

The whole trigger path was tested, from the pixel comparator down to the sector comparator and the camera OR logic. At each position in the trigger chain dedicated scope measurements were performed to assess the level of noise and the purity and inefficiency of the trigger. The measurements showed that the N -majority logic implementation has negligible noise and that for every setting of N exists a threshold setting at which all N -multiplicity trigger are issued and no $N - 1$ -multiplicity signal can trigger.

The only difference between this design and the original one is the fact that the trigger comparator output is sampled, albeit at a very high frequency of 800 MHz, whereas in the old design it was not. This allows for greater flexibility and even for more complex trigger scenarios to be implemented, so it could work as a testbed for concepts put forward for the next generation of Cherenkov telescopes (e.g. digital trigger [9]). With the present scheme, it is expected that the inefficiency inherent to sampling these pulses can be mitigated by lowering the pixel thresholds, while the pixel jitter introduced is negligible (1.25 ns). Extended field studies are pending.

2.3 Ventilation

The new ventilation system, mounted on the back door of the camera, is designed to steadily inject filtered and (optionally) heated air into the camera. It causes a steady overpressure in the entire camera body, and leads to a constant air flow through the drawers and all other small openings of the camera. Like this, a mild homogeneous cooling is achieved, and dust is constantly prevented from entering the camera. The option to heat the incoming air can be used to get rid of humidity after periods of shutdown.

The volume of air exchanged by the ventilation is $> 1300\text{m}^3/\text{hr}$ ($> 3601/\text{s}$), early tests show that such air flow can keep the drawers at the required temperature of $\sim 45^\circ\text{C}$, with a spread of $\pm 5^\circ\text{C}$ depending somewhat on the location of the measurement.

2.4 Network and communication

Drawers, drawer interface box, power distribution box and ventilation system are all connected to the main camera server via Ethernet connections. The camera network has a star topology, all the aforementioned devices communicate only with the central camera server and are independent from one another. The single devices all have 100 Mbit/s copper-based connection to a central switch, which is then connected to the main camera server by means of a 10 Gbit/s optical fiber connection. As a novelty, the camera server is now located remotely in the control building (see Fig. 1), easing its maintenance.

The presence of an ARM-based computer module running Linux on all of these system (with the exception of the ventilation system), makes communication between single components very easy. In order to reduce the slow control and data acquisition software development time, several open source software solutions were adopted.

The operating system running on the Stamp9G45 is a version of Linux, build using the Yocto embedded Linux build system [10], with the kernel module provided by the manufacturer.

Slow control communication is implemented using the Apache Thrift, a lightweight and performant RPC framework [11]. Data transfer was implemented using the ZeroMQ [12] smart socket message-passing library as a transport backbone. Raw event data messages are serialized via an optimized custom protocol, and monitoring data messages are serialized using the google protocol buffer library [13].

These technologies allowed for a robust implementation of network-based slow control and event acquisition to be developed by a very small team of 2 people in short timescales. Evaluation tests showed that the Thrift RPC framework can sustain rates of 10000 requests per second for tens of hours without a single failure, in a busy, complex network. Bulk data transfer tests using the ZeroMQ library have reached link saturation on a 1 GB/s connection, and have exceeded 5 GB/s on a 10 GB/s connection. The latter value is still subject to optimization, and was determined at the time of the test by the CPU speed and memory available rather than by the protocol implementation.

2.5 Production tests

The upgrade requires the production and testing of more than 270 drawers, each hosting 16 analog channels. The drawers are the most complex item of the upgrade, therefore an automatic test bench was specifically set up to ease their testing. An 8-channel pulse generator was designed and

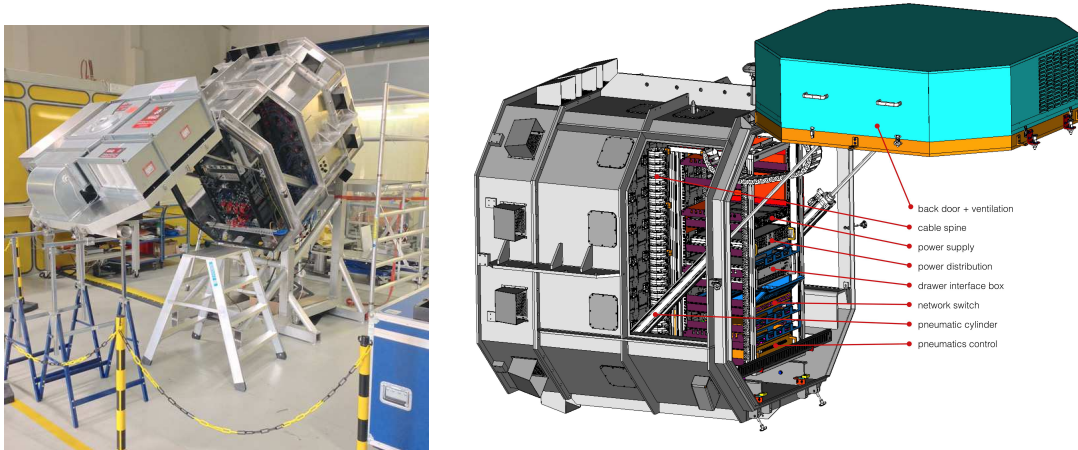


Figure 3: Reproduction of a CT1-4 camera used for test purposes at DESY in Zeuthen. Left: Photo with uncovered back door, such that ventilation components (filters, fans, heating) can be seen. Inside the camera, the rack is mounted with various units installed and blue (trigger) and red (data) cables connecting them. The camera is mounted with the same inclination as in Namibia, to allow for realistic training of deployment. Through the side opening of the camera body, drawer connection boards with their fans can be seen. Right: Labeled sketch of the camera, visualising the layout of the rack units

build for this purpose. Each individual channel on one such pulse generator can deliver a negative pulse with rise and fall times ~ 1 ns, with programmable attenuation up to -54 dB with 1 dB steps (from ~ 300 mV down to ~ 0.6 mV), programmable delay (up to 64 ns in 0.25 ns steps) and programmable width (from 2 ns up to 62 ns in 0.25 ns steps). Furthermore the generator is equipped with a trigger input, a trigger output and a gate, plus a RJ45 connector to mimic the trigger signal protocol of the upgraded camera.

Two of these devices equip the test bench, and permit testing the 16 channels of a drawer simultaneously. More than 300 single unit tests are performed in a test run. The tests are grouped in the following categories: basic functionality, pedestal noise, linearity and cross-talk, and trigger path tests. The pedestal noise and trigger path tests are performed by the ARM module on drawer itself, since they do not require any external input. Therefore, it will be possible to test most of the features of a drawer even after its installation on the camera.

3. Integration test bench

In order to test the interplay of the new components as good as possible before deployment, a replica of a H.E.S.S. I camera was produced and assembled in Zeuthen (Fig. 3). Even without the photomultiplier tubes (PMTs) available, the trigger chain, camera-internal network, cable mapping, ventilation, slow control, power supply and mechanical integration could thus be tested to a good extent.

An additional dark chamber with a 4-drawer “mini camera”, including PMTs, can be attached to the test camera to create a realistic dark trigger setup. A copy of the H.E.S.S. data acquisition system (DAQ) was attached and the original state transitions of the system were emulated (see [5], Fig. 3).

4. Outlook and future perspective

The preliminary plan is to install the first camera on CT1 in 2015, and the remaining three in one campaign in 2016, after an extensive verification of the CT1 camera. Besides the improved dead time, the design of the H.E.S.S. I Upgrade opens many possibilities to optimise the trigger logic, front-end readout window and time information of pulses. The upgraded H.E.S.S. cameras can therefore be expected to improve the sensitivity of H.E.S.S., but also continue to provide a test bed for various technical innovations that are being discussed in the framework of CTA R&D efforts.

Acknowledgements. The support of the Namibian authorities and of the University of Namibia in facilitating the construction and operation of H.E.S.S. is gratefully acknowledged, as is the support by the German Ministry for Education and Research (BMBF), the Max Planck Society, the German Research Foundation (DFG), the French Ministry for Research, the CNRS-IN2P3 and the Astroparticle Interdisciplinary Programme of the CNRS, the U.K. Science and Technology Facilities Council (STFC), the IPNP of the Charles University, the Czech Science Foundation, the Polish Ministry of Science and Higher Education, the South African Department of Science and Technology and National Research Foundation, and by the University of Namibia. We appreciate the excellent work of the technical support staff in Berlin, Durham, Hamburg, Heidelberg, Palaiseau, Paris, Saclay, and in Namibia in the construction and operation of the equipment.

References

- [1] H.E.S.S. Collaboration, *Observations of the Crab nebula with HESS*, A&A 457 (2006) 899
- [2] P. Vincent et al., *Performance of the H.E.S.S. cameras* Proc. 28th ICRC (2003)
- [3] C. L. Naumann et al., *New electronics for the Cherenkov Telescope Array (NECTAr)* NIM A 695 (2012) 44
- [4] E. Delagnes, *Specifications of the Nectar0 Chip v.3*, Internal Document, Irfu CEA Saclay (2011)
- [5] A. Balzer et al., *The H.E.S.S. central data acquisition system* Astropart. Phys. 54 (2014) 67
- [6] MAGIC Collaboration, *The major upgrade of the MAGIC telescopes, Part I: The hardware improvements and the commissioning of the system* acc. for publication in Astropart. Phys. (2015) arXiv:1409.6073
- [7] D. B. Kieda for the VERITAS Collaboration *The Gamma Ray Detection sensitivity of the upgraded VERITAS Observatory*, Proc. 33rd ICRC (2013)
- [8] Funk, S. et al. *The trigger system of the H.E.S.S. telescope array*. Astropart. Phys. 22, 285–296 (2004).
- [9] U. Schwanke et al., *A versatile digital camera trigger for telescopes in the Cherenkov Telescope Array*. Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment 782, 92–103 (2015).
- [10] *Yocto Project | Open Source embedded Linux build system, package metadata and SDK generator*, <https://www.yoctoproject.org/> (retrieved on 2015-06-25).
- [11] *Apache Thrift - Home*, <https://thrift.apache.org/> (retrieved on 2015-06-25).
- [12] *Code Connected - zeromq*, <http://zeromq.org/> (retrieved on 2015-06-25).
- [13] *Protocol Buffers.*, <https://developers.google.com/protocol-buffers> (retrieved on 2015-06-25)