



HAL
open science

Galactic tides and the Crater II dwarf spheroidal: a challenge to LCDM?

Alexandra Borukhovetskaya, Julio Navarro, Raphaël Errani, Azadeh Fattahi

► **To cite this version:**

Alexandra Borukhovetskaya, Julio Navarro, Raphaël Errani, Azadeh Fattahi. Galactic tides and the Crater II dwarf spheroidal: a challenge to LCDM?. Monthly Notices of the Royal Astronomical Society, 2022, 10.1093/mnras/stac653 . hal-03615072

HAL Id: hal-03615072

<https://hal.science/hal-03615072v1>

Submitted on 31 Mar 2023

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Galactic tides and the Crater II dwarf spheroidal: a challenge to LCDM?

Alexandra Borukhovetskaya ¹, ¹★ Julio F. Navarro,¹ Raphaël Errani^{1,2} and Azadeh Fattahi ³

¹Department of Physics and Astronomy, University of Victoria, Victoria, BC V8P 5C2, Canada

²Universite de Strasbourg, CNRS, Observatoire astronomique de Strasbourg, UMR 7550, F-67000 Strasbourg, France

³Institute for Computational Cosmology, Department of Physics, University of Durham, South Road, Durham DH1 3LE, UK

Accepted 2022 March 2. Received 2022 March 1; in original form 2021 December 2

ABSTRACT

The unusually low velocity dispersion and large size of Crater II pose a challenge to our understanding of dwarf galaxies in the Lambda cold dark matter (LCDM) cosmogony. The low velocity dispersion suggests either a dark halo mass much lower than the minimum expected from hydrogen cooling limit arguments or one that is in the late stages of extreme tidal stripping. The tidal interpretation has been favoured in recent work and is supported by the small pericentric distances consistent with available kinematic estimates. We use N -body simulations to examine this interpretation in detail, assuming a Navarro–Frenk–White (NFW) profile for Crater II’s progenitor halo. Our main finding is that, although the low velocity dispersion can indeed result from the effect of tides, the large size of Crater II is inconsistent with this hypothesis. This is because galaxies stripped to match the observed velocity dispersion are also reduced to sizes much smaller than the observed half-light radius of Crater II. Unless its size has been substantially overestimated, reconciling this system with LCDM requires that either (i) it is not bound and near equilibrium (unlikely, given its crossing time is shorter than the time elapsed since pericentre) or (ii) its progenitor halo deviates from the assumed NFW profile. The latter alternative may signal that baryons can affect the inner halo cusp even in extremely faint dwarfs or, more intriguingly, may signal effects associated with the intimate nature of the dark matter, such as finite self-interactions, or other such deviations from the canonical LCDM paradigm.

Key words: galaxies: dwarf – galaxies: evolution – dark matter.

1 INTRODUCTION

The Crater II dwarf spheroidal (dSph) is a distant Milky Way (MW) satellite discovered in imaging data from the ATLAS survey at the VLT Survey Telescope by Torrealba et al. (2016). Its unusual properties were immediately apparent. Although its total luminosity ($M_V \sim -8$) is comparable to that of the faintest ‘classical’ dSphs, such as Draco and Ursa Minor, Crater II’s enormous size (projected half-light radius, $R_{1/2} \sim 1$ kpc, and spanning nearly 5° across the sky) is comparable to that of Fornax, a dSph more than 100 times more luminous. Indeed, Crater II is one of the lowest surface brightness galaxies ever discovered, several decades fainter than the ultradiffuse galaxy population identified by surveys such as Dragonfly (van Dokkum et al. 2015).

These unusual photometric properties are compounded by equally unusual kinematics. The most recent estimates put the velocity dispersion of its stars at $\sigma_{\text{los}} \approx 2.3 \text{ km s}^{-1}$, one of the lowest amongst all known dSphs but still large enough to imply the presence of large amounts of dark matter ($M/L \approx 50$, Caldwell et al. 2017; Fu, Simon & Alarcón Jara 2019; $M/L \approx 30$, Ji et al. 2021). Such uncommon properties point to atypical formation paths. Suggestions include the possibility that Crater II formed in a halo of unusually low density (Amorisco 2019), or that its structure was severely affected by Galactic tides after formation, or both.

The latter explanation is currently the most popular. Frings et al. (2017) and Applebaum et al. (2021), for example, report simulations

of satellites whose properties are similar to those of Crater II after undergoing severe tidal stripping. Similar conclusions were reported by Sanders, Evans & Dehnen (2018), who cautioned, however, that it was difficult to account simultaneously for the velocity dispersion and size of Crater II in cuspy dark matter haloes such as those expected in the Lambda cold dark matter (LCDM) cosmogony (Navarro, Frenk & White 1996b, 1997). Fattahi et al. (2018) extrapolated the ‘tidal tracks’ of Peñarrubia, Navarro & McConnachie (2008) and Errani, Peñarrubia & Tormen (2015) to reach a similar conclusion, noting that explaining Crater II through tidal stripping implied the rather extreme case of a progenitor that had lost more than 99 per cent of its initial stellar content to tides.

Despite these difficulties, consensus for a tidal interpretation of the unusual properties of Crater II (and for those of Antlia II, another ‘feeble giant’ MW satellite; Torrealba et al. 2019) seems to have emerged. In the case of Antlia II, the case for tides has been strengthened by the detection of a clear velocity gradient aligned with the orbital path indicated by the latest proper motions from *Gaia* (Ji et al. 2021).

No such gradient, however, is clearly present in Crater II, at least over the area surveyed spectroscopically. This is perhaps not surprising: Crater II’s negative Galactocentric radial velocity suggests that it is at present just past apocentre; together with its large distance (~ 120 kpc), this implies that its most recent pericentric passage must have occurred about ~ 1 Gyr ago. As discussed¹ by Peñarrubia et al. (2009), the inner regions of a bound tidal remnant relax quickly, pushing all signatures of tidal disturbance to regions where the local

* E-mail: asya@uvic.ca

¹See also earlier work by Aguilar & White (1986) and Navarro (1990).

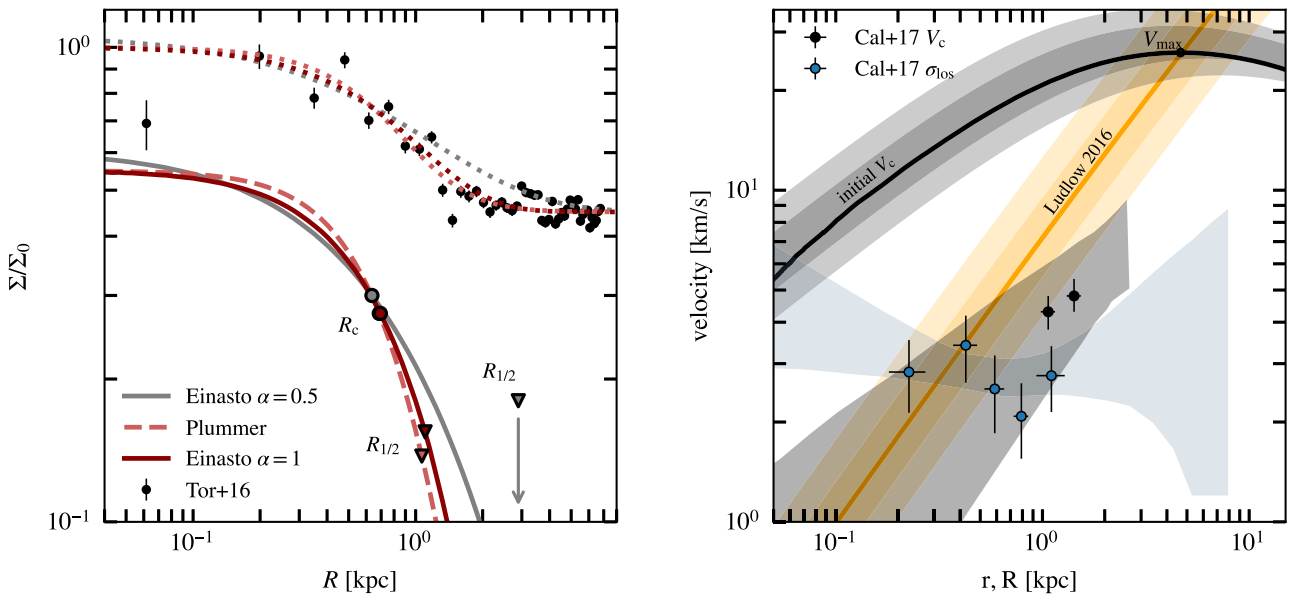


Figure 1. Observed structural parameters of Crater II. The left-hand panel shows the observed surface brightness taken from Torrealba et al. (2016) together with fits adopting different models for the density profile of the stellar component. Light red curves correspond to a Plummer (1911) fit with the parameters of Torrealba et al. (2016). Dotted profiles include the contribution of a constant background level equal to $0.45 \times$ the central surface brightness, Σ_0 . Dark red curves correspond to an $\alpha = 1$ Einasto profile (equation 2), which fits the Plummer projected density profile quite well. Grey curves show an $\alpha = 0.5$ Einasto profile with similar core radius, R_c (circles), but a much larger half-light radius ($R_{1/2}$, triangles) than the other profiles. All seem to provide acceptable fits to the data, implying that, unlike R_c , Crater II’s $R_{1/2}$ is not well constrained. The right-hand panel shows the measured velocity dispersion of Crater II (blue circles and shaded area) as well as corresponding estimates for the circular velocity, taken from Caldwell et al. (2017). The black line (and shaded area) shows the circular velocity profile of our Crater II NFW model before stripping, motivated by results from the APOSTLE simulation for galaxies of similar stellar mass (Fattahi et al. 2018). The yellow line shows the r_{\max} - V_{\max} relation for LCDM haloes at $z = 0$ from Ludlow et al. (2016). The inner shaded bands correspond to $\pm 1\sigma$ scatter in concentration and the outer fainter bands to $\pm 2\sigma$.

crossing time exceeds the time elapsed since pericentre, $t - t_p$ (see their equation 5). Assuming $(t - t_p) \sim 1$ Gyr, and using the measured velocity dispersion, this implies that the inner $R \lesssim 1.3$ kpc should be close to dynamical equilibrium. This region is similar to the area over which radial velocities are available (~ 0.65 from the dwarf’s centre; Ji et al. 2021).

This insight simplifies considerably the analysis, as it disfavors interpretations where the singular properties of Crater II are due to large, transient departures from equilibrium. If Galactic tides have indeed been responsible for shaping Crater II, then its properties must be consistent with the structure of the remnants of tidally stripped LCDM subhaloes, an issue that has been studied extensively over the years using N -body simulations (see e.g. Hayashi et al. 2003; Peñarrubia et al. 2008; Errani et al. 2015). The latest work suggests that LCDM subhaloes, if well approximated by cuspy Navarro-Frenk-White profiles (hereafter NFW, see Navarro, Frenk & White 1996b, 1997), should almost always leave behind a bound remnant (Peñarrubia et al. 2010; van den Bosch & Ogiya 2018; van den Bosch et al. 2018). The remnant properties are fully specified by the initial characteristic crossing time of the subhalo, by the number of orbits completed, and by the orbital time at pericentre (Errani & Navarro 2021).

This discussion thus calls into question whether Crater II is actually a stellar system inhabiting a subhalo nearing full tidal disruption. Are the peculiar size and kinematics of Crater II actually consistent with the tidal remnants of LCDM haloes?

We address this issue here using N -body simulations of the tidal evolution of NFW subhaloes in orbits chosen to match the observed present-day position and velocity of Crater II in the MW halo potential. We choose cuspy NFW haloes for this work not only

because this is one of the best tested LCDM predictions, but also because the halo inner structure of galaxies as faint as Crater II should be relatively unaffected by the effects of baryons, which could in principle alter the inner density cusp in more massive galaxies (see e.g. Navarro, Eke & Frenk 1996a; Read & Gilmore 2005; Pontzen & Governato 2012; Di Cintio et al. 2014, and references therein).

This paper is organized as follows. We first summarize the observed properties of Crater II, as well as our numerical setup, in Section 2. The results of the simulations, for both dark matter and stellar tracers, are analysed and presented in Section 3. We end with a brief summary and a discussion of our results in a cosmological context in Section 4.

2 OBSERVATIONS AND SIMULATIONS

This section summarizes the observed properties of Crater II, and outlines the numerical setup of the N -body simulations used to model the tidal evolution of Crater II in the gravitational potential of the MW.

2.1 Observed properties of Crater II

The photometric properties of Crater II were estimated by Torrealba et al. (2016) in its discovery paper. These authors estimate a projected half-light radius of the order of $R_{1/2} \sim 1.066$ kpc by fitting a Plummer model to star counts. We reproduce their published data in the left-hand panel of Fig. 1, which shows that the number density profile spans only a factor of ~ 2 in density before becoming dominated by foreground stars. This results in large uncertainties in estimates of the half-light radius, depending on the assumed profile shape. This is

Table 1. Parameters of the analytical, static MW potential used in this study. The model is a re-parametrization of the McMillan (2011) model, as discussed in Errani & Peñarrubia (2020).

Component	Functional form			
Disc (thin)	Miyamoto & Nagai (1975)	$M = 5.9 \times 10^{10} M_{\odot}$	$a_d = 3.9$ kpc	$b_d = 0.31$ kpc
Disc (thick)	Miyamoto & Nagai (1975)	$M = 2.0 \times 10^{10} M_{\odot}$	$a_d = 4.4$ kpc	$b_d = 0.92$ kpc
Bulge	Hernquist (1990)	$M = 2.1 \times 10^{10} M_{\odot}$	$a = 1.3$ kpc	
DM halo	Navarro et al. (1997)	$M_{200} = 1.15 \times 10^{12} M_{\odot}$	$r_s = 20.2$ kpc	$c = r_{200}/r_s = 9.5$

Table 2. Current observational constraints, as well as parameters of the three orbits explored using N -body simulations in this study. The Fritz et al. (2018) proper motions are followed first by the statistical error and second by the systematic error. Orbits 1 and 2 are the orbits corresponding to the 16th and 84th percentile of the distribution in pericentres obtained with the proper motions of Kallivayalil et al. (2018). Orbit 3 is the orbit corresponding to the median observed quantities of McConnachie & Venn (2020). Pericentres and apocentres are computed for the MW potential model discussed in Section 2.2.

Observation	α	δ	Distance (kpc)	μ_{α^*} (mas yr $^{-1}$)	μ_{δ} (mas yr $^{-1}$)	v_r (km s $^{-1}$)
	$11^{\text{h}}49^{\text{m}}12^{\text{s}}{}^a$	$-18^{\circ}24'0''{}^a$	$117.5 \pm 1.1{}^a$	$-0.246 \pm 0.052{}^b$ $-0.184 \pm 0.061 \pm 0.035{}^c$ $-0.07 \pm 0.02{}^d$	$-0.227 \pm 0.026{}^b$ $-0.106 \pm 0.031 \pm 0.035{}^c$ $-0.11 \pm 0.01{}^d$	$87.5 \pm 0.4{}^e$
a Torrealba et al. (2016), b Kallivayalil et al. (2018), c Fritz et al. (2018), d McConnachie & Venn (2020), e Caldwell et al. (2017)						
Model parameters	Pericentre (kpc)	Apocentre (kpc)	Distance (kpc)	μ_{α^*} (mas yr $^{-1}$)	μ_{δ} (mas yr $^{-1}$)	v_r (km s $^{-1}$)
Orbit 1	4.23	130	116	-0.169	-0.267	87.8
Orbit 2	15.5	133	117	-0.102	-0.225	87.2
Orbit 3	37.4	139	118	-0.07	-0.11	87.5

illustrated by the grey and dark red lines in Fig. 1, which both fit the observations fairly well, although their half-light radii (as marked by triangles) differ by a factor of ~ 3 . A much better constrained parameter is the core radius, $R_c \approx 700$ pc, defined as the radius where the surface brightness of a galaxy drops by a factor of 2 from the central value.

Stellar half-light radii are also difficult to measure in numerical simulations, particularly in the cases where tides have led to substantial amounts of mass-loss. In these cases, estimated $R_{1/2}$ values depend critically on which stars are included in the analysis. Including weakly bound, or escaping, stars, for example, typically results in poorly defined estimates vulnerable to sizeable transient fluctuations. Because of this, we shall adopt R_c (marked with circles in the left-hand panel of Fig. 1) rather than $R_{1/2}$ (triangles) as the characteristic photometric radius of Crater II, although our main conclusions do not depend critically on this choice. For a Plummer model, which is often adopted in observational studies to fit the stellar density profile, $R_c \approx 0.6 R_{1/2}$.

The right-hand panel of Fig. 1 summarizes some of the kinematic information available for Crater II, compiled from Caldwell et al. (2017). [We shall hereafter use lower-case r to denote three-dimensional (3D) radii and upper-case R for projected radii.] The observed velocity dispersion profile is shown in grey-blue. The black circles indicate corresponding circular velocity estimates obtained using the Wolf et al. (2010) and Walker et al. (2009) mass estimators.²

For comparison, we show with a thick black curve the (NFW) circular velocity profile expected for an *isolated* dwarf galaxy of stellar mass comparable to that of Crater II (assumed to be $M_{\star} \approx 2.56 \times 10^5 M_{\odot}$, see Section 2.4 for details). According to the results of the APOSTLE (A Project Of Simulating The Local

Environment) cosmological hydrodynamical simulations (Fattahi et al. 2018), galaxies like Crater II are expected to form in haloes with peak circular velocity $V_{\text{max}} \sim 26$ km s $^{-1}$, or a virial³ mass $M_{200} \sim 2.7 \times 10^9 M_{\odot}$. For such model to be viable, it is clear from Fig. 1 that Galactic tides must have led to a large depletion of the original dark matter content of Crater II.

2.2 MW potential model

The MW host is represented by an analytical, static potential, which is comprised of an axisymmetric two-component Miyamoto & Nagai (1975) disc, a Hernquist (1990) bulge, and an NFW dark matter halo. The model parameters are taken from Errani & Peñarrubia (2020) – i.e. chosen to approximate the McMillan (2011) model, wherein at the solar circle $R_0 = 8.29$ kpc, the circular velocity is $V_c = 240$ km s $^{-1}$. The thick and thin Miyamoto & Nagai (1975) discs are each parametrized by a disc mass M , a scale length a_d , and a scale height b_d . The Hernquist (1990) bulge is similarly defined with a total mass M and scale length a . Finally, the MW dark matter halo can be characterized by an NFW halo with a scale radius r_s and a virial mass M_{200} . The host parameters are summarized in Table 1.

2.3 Orbits

The orbit of Crater II in the assumed MW potential may be estimated from its present-day Galactocentric position and velocity, as inferred from its sky position, radial velocity, distance, and proper motion, assuming that the effects of dynamical friction can be neglected. Of

²These mass estimates are supported by the full Jeans modelling of Caldwell et al. (2017) (shaded region).

³We define virial quantities as those within a sphere of mean density equal to 200 times the critical density for closure, $\rho_{\text{crit}} = 3H_0^2/8\pi G$. We use the subscript ‘200’ to indicate virial quantities.

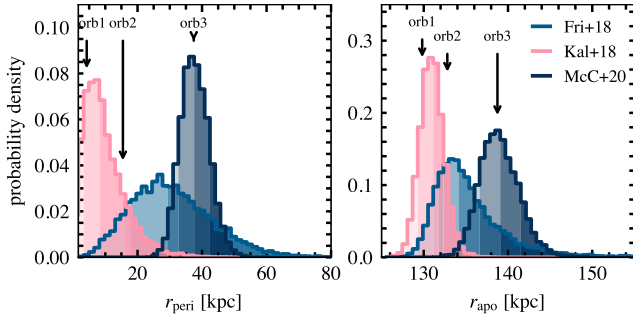


Figure 2. The probability distributions of Crater II pericentric (left-hand panel) and apocentric (right-hand panel) distances obtained from the proper motions, distance, and radial velocity, together with their reported uncertainties. Each colour corresponds to different proper motion estimates, from Kallivayalil et al. (2018, pink), Fritz et al. (2018, blue), and McConnachie & Venn (2020, black). The three orbits selected to explore the available parameter space are shown with arrows pointing down. Regions beneath the histograms darken at the 16th, 50th, and 84th percentile.

the observed parameters (summarized in Table 2), the proper motions contribute the majority of the uncertainty budget.

Assuming $\mu_{\alpha^*} = -0.246 \pm 0.052$ mas yr⁻¹, $\mu_{\delta} = -0.227 \pm 0.026$ mas yr⁻¹ (Kallivayalil et al. 2018),⁴ for example, these uncertainties result in a fairly broad distribution of possible pericentric and apocentric distances, shown by the pink histograms in Fig. 2. Different estimates for the proper motion will, of course, yield different orbits. We show also in Fig. 2 the pericentric and apocentric distributions corresponding to proper motion estimates from McConnachie & Venn (2020) and Fritz et al. (2018), illustrating that a wide range of orbits are permissible given the current data.

Because of the large allowed range, we have explored three different orbits, varying mainly the orbital pericentric distance (the most critical parameter for tidal effects), $r_{\text{peri}} = 4.2, 16,$ and 37 kpc. The corresponding orbits are hereafter referred to as orbits 1, 2, and 3, respectively, with parameters summarized in Table 2. Initial conditions for the orbits of N -body runs are obtained by integrating them backwards in time for 10 Gyr. The shapes of the resulting orbits are shown in Fig. A1 in a coordinate system where the Sun is located at $(X, Y, Z)_{\odot} = (-8.3, 0, 0)$ kpc, and the velocity of the local standard of rest is in the positive Y direction.

2.4 Crater II model

2.4.1 Dark matter component

The Crater II halo is modelled as an equilibrium N -body realization of the NFW density profile,

$$\rho_{\text{NFW}}(r) = \frac{M_{200}}{4\pi r_s^3} \frac{(r/r_s)^{-1}(1+r/r_s)^{-2}}{[\ln(1+c) - c/(1+c)]}. \quad (1)$$

This profile is fully specified by two parameters: e.g. a virial mass, M_{200} , and concentration, $c = r_{200}/r_s$, or, alternatively, by a maximum circular velocity, V_{max} , and the radius at which it is reached, r_{max} .

Cosmological hydrodynamical simulations, such as the APOSTLE suite of Local Group simulations (Fattahi et al. 2016; Sawala et al. 2016), have shown that V_{max} correlates strongly with galaxy stellar

⁴Where with μ_{α^*} we designate the proper motion in α including the $\cos(\delta)$ factor.

Table 3. Current properties of Crater II (M_* , $V_{1/2}$, $r_{1/2}$) and inferred structural parameters at infall of the dark matter halo (V_{max} , r_{max}) and stellar component (α , R_{core} , M_*). Half-light radius $r_{1/2}$ and circular velocity $V_{1/2}$ at the half-light radius are as in Caldwell et al. (2017), taking $r_{1/2} = 4/3R_{1/2}$.

Observation	M_* ($10^5 M_{\odot}$)	$V_{1/2}$ (km s ⁻¹)	$r_{1/2}$ (kpc)
Initial halo	2.56	$4.8^{+0.6}_{-0.5}$	1.421 ± 0.112
	V_{max} (km s ⁻¹)	r_{max} (kpc)	
	25.9	4.70	
Initial stars	α	R_{core} (kpc)	
E1	1	0.46	
E2	1	0.91	
E3	1	1.79	

mass, M_* (see also Santos-Santos et al. 2021, for a recent compilation). We adopt a stellar mass for Crater II of $M_* = 2.56 \times 10^5 M_{\odot}$, computed from the absolute magnitude of Torrealba et al. (2016) and assuming $M_*/L_V = 1.6$, typical of dwarf galaxies in the Local Group (Woo, Courteau & Dekel 2008).

As in Borukhovetskaya et al. (2021), we estimate V_{max} (before tidal effects) using the empirical fit⁵ to the $M_* - V_{\text{max}}$ correlation in APOSTLE from Fattahi et al. (2018). The characteristic radius r_{max} then follows from the Ludlow et al. (2016) parametrization of the LCDM halo mass–concentration relation at redshift $z = 0$. As listed in Table 3, the resulting NFW profile has $V_{\text{max}} = 25.9$ km s⁻¹ and $r_{\text{max}} = 4.7$ kpc, or, in terms of virial mass and concentration, $M_{200} = 2.72 \times 10^9 M_{\odot}$ and $c = 13.6$. The corresponding circular velocity profile is shown in the right-hand panel of Fig. 1.

2.4.2 Stellar component

The stellar component of Crater II is modelled assuming that it contributes negligibly to the total gravitational potential. We consider, in particular, stellar components modelled as Einasto profiles (Einasto 1965),

$$\rho_E(r) = \rho_{E0} \exp[-(r/r_E)^\alpha], \quad (2)$$

with $\alpha = 1$ (i.e. exponential spheres) and three different core radii: $R_c = 0.5, 0.9,$ and 1.7 kpc. These values have been chosen to bracket the observed present-day core radius of Crater II, which is ~ 0.7 kpc. These three stellar models are hereafter referred to as E1, E2, and E3, respectively, and their parameters are summarized in Table 3. Core radii are related to the Einasto scale radius, r_E , by the relation $R_c \approx 1.24 r_E$. In practice, stars are modelled⁶ by attaching a probability to each dark matter particle, following the appropriate distribution function (see e.g. Bullock & Johnston 2005). Using this approach, the probability of each N -body particle to represent a star relates the dark matter and stellar energy distributions as

$$P(E) \propto (dN/dE)_*/(dN/dE). \quad (3)$$

The underlying energy distributions are computed numerically as outlined in Errani & Peñarrubia (2020). These probabilities are

⁵ $M_* = M_0 v^\alpha \exp(-v^\gamma)$, where $v = V_{\text{max}}/50$ km s⁻¹, and $(M_0, \alpha, \gamma) = (3.0 \times 10^8 M_{\odot}, 3.36, -2.4)$.

⁶For details, see the publicly available implementation of Errani & Peñarrubia (2020): <https://github.com/rerrani/nbody>

computed at infall and followed throughout the orbit, where at any point the stellar structure and kinematics may be recovered from the dark matter distribution by applying the individual stellar probabilities as weights.

2.5 Simulation code

We use 10^7 -particle N -body realizations of NFW haloes, computed using the `ZENO`⁷ software package developed by Joshua Barnes at the University of Hawaii. This software uses Monte Carlo sampling of a given distribution function to generate systems in virial equilibrium.

We allow the halo to fully relax prior to introducing it into the MW potential, by running it first in isolation for 5 Gyr using the publicly available `GADGET-2` simulation code (Springel 2005). Once the halo has relaxed, the N -body model is evolved on each of the orbits described in Section 2.3 and in the potential of Section 2.2 for ~ 10 Gyr. Forces between particles are smoothed with a Plummer-equivalent softening length of $\epsilon_p = 7$ pc, and we consider our results converged outside $r_{\text{conv}} = 84$ pc.⁸

3 RESULTS

Having introduced our numerical setup, we examine next the tidal evolution of Crater II. We discuss first the evolution of the dark matter (Section 3.1), before discussing the evolution of embedded stellar components in Section 3.2.

3.1 Tidal effects on the dark matter component

The evolution of the circular velocity within $R_{1/2} = 1.066$ kpc of our Crater II halo models is shown in Fig. 3. Different colours correspond to different orbits. A horizontal black line indicates, for reference, the estimated circular velocity from Caldwell et al. (2017).

As tides strip the system, the circular velocity decreases continuously, with abrupt drops corresponding to subsequent pericentric passages. The magnitude of the decrease is heavily dependent on the assumed pericentric distance. We see that for a pericentre as large as ~ 40 kpc (orbit 3), tidal effects are not enough to reduce the circular velocity enough to match the observational estimate.

The mass-loss on orbits 1 and 2, with pericentres of 4 and 16 kpc, respectively, seems large enough to bring the initial halo into agreement with the observed estimate in less than 10 Gyr. This happens after only two pericentric passages for orbit 1, but it takes five full orbits for orbit 2. We note that bringing the assumed halo into agreement with the observed V_c implies a dramatic amount of mass-loss: only 0.2 per cent of the initial dark mass remains bound when $V_c(R_{1/2})$ approaches the observed value of ~ 4 km s⁻¹.

We conclude that it is in principle possible to explain the unusually low velocity of Crater II if the dSph is placed on orbits with pericentric distances of the order of ~ 15 kpc or less. This value is consistent with current observational estimates, though tighter constraints on the allowed pericentres should be able to rule out this scenario if larger pericentres are found to be favoured.

3.2 Tidal effects on the stellar component

The stellar components are also affected by tidal losses, to an extent that depends on the assumed initial density profile and radial

⁷<http://www.ifa.hawaii.edu/faculty/barnes/zeno/>

⁸We define the radius of convergence, r_{conv} , as the innermost radius where the initial circular velocities deviate by less than ~ 1 per cent from the target NFW profile.

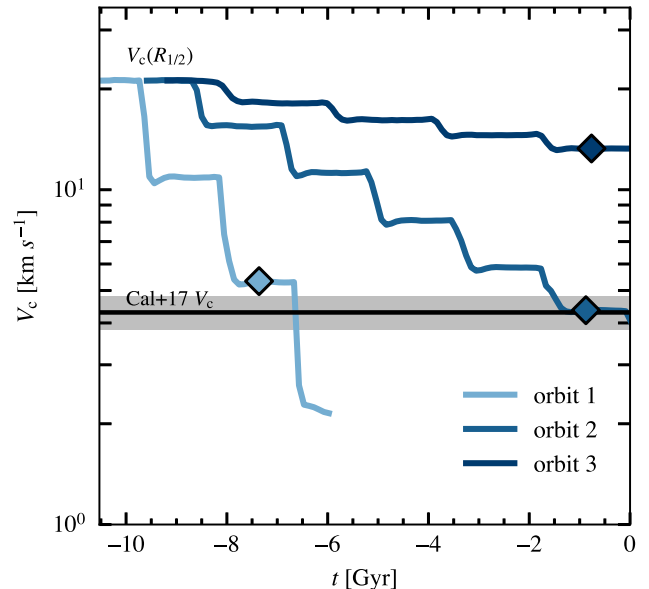


Figure 3. Evolution of the circular velocity at the estimated half-light radius of Crater II (i.e. V_c at $r = R_{1/2} = 1.066$ kpc) for our halo model in each of the three different orbits explored in this work; lighter shades correspond to smaller pericentres. The horizontal black solid line and grey shaded region represent the observational constraint on V_c , as well as the $\pm 1\sigma$ uncertainty interval, taken from Caldwell et al. (2017). Diamonds indicate the snapshot(s) we identify for comparison with observational data. Note that, on orbit 3, our chosen halo does not get stripped enough to bring the system into agreement with the observed constraint after orbiting for ~ 10 Gyr.

segregation of stars relative to the dark matter (see Errani et al. 2021, for a more detailed discussion). We focus here on models E1, E2, and E3 – exponential spheres that differ mainly in their initial core/half-light radii (see Section 2.4 for details). We limit our analysis to the evolution on orbit 2, but very similar results are obtained for orbit 1, albeit on a compressed time-scale.

We begin by tracking the evolution of the core radius, R_c , and of the line-of-sight velocity dispersion, σ_{los} , of the stars (averaged within R_c) in Fig. 4. Results for E1, E2, and E3 are shown with different coloured tracks, starting with the initial conditions (triangles) and ending, after five pericentric passages, on the coloured circles highlighted in black. Each small circle along the tracks indicates a subsequent apocentric passage. The ‘target’ Crater II core radius and velocity dispersion inside the core radius are shown by a red circle with a blue error bar.

Interestingly, none of the stellar tracks seems to approach the observed location of Crater II in the R_c – σ_{los} plane. As the system loses mass, the velocity dispersions decrease, but so do the core radii of the stars. By the time the velocities approach the observed value, the core radii are almost four times smaller than observed. This result applies to all three stellar models, regardless of their initial radius.

Because of this discrepancy, reproducing the observed surface density profile of stars in Crater II is not possible, as shown in the left-hand panel of Fig. 5. The solid curves in that figure indicate the final density profile of bound stars in orbit 2, normalized to match approximately the central observed value for Crater II with background. Dotted curves in Fig. 5 show the same profiles after adding a constant background of $\Sigma_{\text{bck}} = 0.45 \Sigma_0$. The bound stellar remnant has a similar size in all cases, and is clearly too small to match the observed profile for Crater II.

This result is general, and not just particular to our choices of density profile, or initial radial segregation. This may be seen by noticing

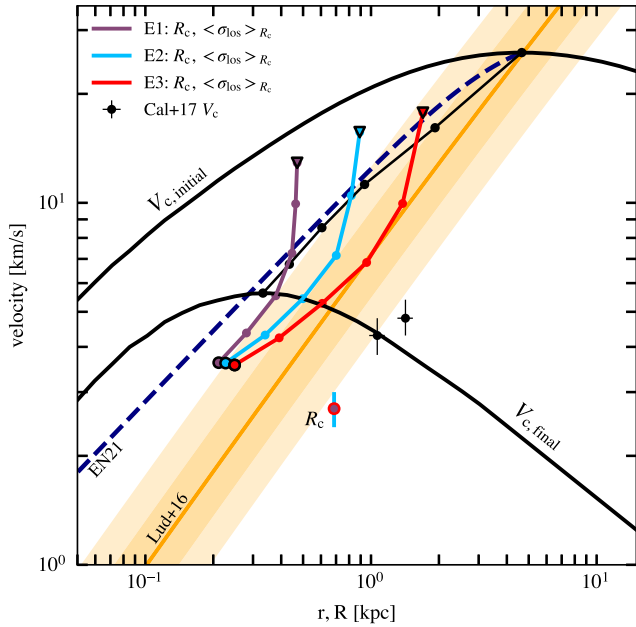


Figure 4. Evolution of the core radius, R_c , and line-of-sight velocity dispersion, σ_{los} , for three different stellar component models (E1, E2, and E3) on orbit 2. Triangles indicate the initial values, and coloured circles the values at the final snapshot. The coloured lines trace the evolution of these parameters as a function of time; each small circle marks subsequent apocentric passages. The orange line shows the $r_{\text{max}}-V_{\text{max}}$ relation (plus 1σ and 2σ scatter bands in concentration) expected for LCDM haloes at $z=0$ (Ludlow et al. 2016, in orange). The initial and final circular velocities of the assumed Crater II progenitor halo are shown by the solid black curves. Black circles with error bars show the observed V_c constraint from Caldwell et al. (2017). The blue-red circle indicates the observed values of R_c and σ_{los} for Crater II. The dashed curve shows the ‘tidal track’ of Errani & Navarro (2021), which traces quite well the evolution of r_{max} and V_{max} of the halo (shown as a solid black line with circles at each apocentric passage) as a function of time.

that, after substantial mass-loss, the stellar tracks of all three systems follow the tidal track of the remnant dark matter component, indicated by the dashed black line in Fig. 4. This track follows the evolution of the remnant’s r_{max} and V_{max} characteristic parameters, which, by the end of the simulation on orbit 2, has seen its V_{max} decrease by a factor of 5 and its r_{max} by a factor of 14. More importantly, as discussed by Errani & Navarro (2021), the remnant subhalo approaches a mass profile that is well approximated by an ‘exponentially truncated cusp’, with r_{max} roughly delineating a ‘tidal truncation’ radius.

The tidal truncation is easy to spot in the right-hand panel of Fig. 5, where we plot the line-of-sight velocity dispersion profiles of the bound stellar remnants at various times during the evolution. Velocity dispersions are roughly constant in the inner regions, but drop sharply outside r_{max} , the location of which is marked with an ‘X’ symbol on each curve. The decline is reversed and the velocity dispersion rises in the outermost parts because of the presence of weakly bound stars moving radially outwards. The position of this upturn approximately coincides with the location where the local crossing time equals the time elapsed since the previous pericentric passage, as discussed by Peñarrubia et al. (2009). The locations of this ‘break radius’, computed⁹ as $R_b = C \sigma_0 (t - t_p)$ (where $C = 0.35$ and σ_0 is the central line-of-sight velocity dispersion computed as

⁹Note that the value of the constant $C = 0.35$ used here differs from the $C = 0.55$ value of equation (5) in Peñarrubia et al. (2009) because in that paper R_b

the average within $1/2 r_{\text{max}}$), are indicated by small squares on each of the curves shown in the right-hand panel of Fig. 5.

In other words, in these late stages of tidal stripping, the size of the stellar component cannot exceed r_{max} because there is little bound mass left outside r_{max} (see, also, Kravtsov 2010; Errani & Navarro 2021). Indeed, an exponential cusp profile has only ~ 13 per cent of mass outside $2r_{\text{max}}$, and just ~ 3 per cent outside $3r_{\text{max}}$. This also readily explains why all three stellar models converge to a remnant with the same R_c and σ_{los} , regardless of their initial size. We emphasize that this is a general result; the results for other orbits are shown in Fig. B1 and the results for an alternate density profile in Fig. B2. We discuss next our results in the context of earlier work, as well as the implications of our findings for the interpretation of satellites like Crater II in a cosmological context.

3.3 Comparison with earlier work

The findings discussed above seem to contradict the results of Frings et al. (2017) and of Applebaum et al. (2021), who used simulations of tidally stripped satellites in a cosmological context to argue that it is possible to form Crater II-like satellites in LCDM cosmological simulations. We examine the reasons for the disagreement next, but emphasize that the discussion of their work relates *only* to their claims regarding Crater II-like systems. This discussion is not meant as wholesale criticism of their work, but rather as seeking a plausible explanation of why our conclusions differ from theirs on this particular topic.

We begin our discussion with an analysis of the results reported by Frings et al. (2017). Their ‘satellites I and II’, in particular, seem to increase substantially in size as their haloes are tidally stripped, at odds with our results. This ‘tidal expansion’, however, occurs only when the orbit chosen leads to extreme tidal stripping (orbit V in the notation of Frings et al. 2017).

Satellite I inhabits a halo with a pronounced ‘core’ in the density profile, and thus differs from the NFW halo models studied here (Macciò et al. 2017). For ‘cored’ haloes, tidal expansion of its stellar component under extreme tidal stripping is actually expected (see e.g. Errani et al. 2015; Sanders et al. 2018). Furthermore, satellite I is a fairly massive satellite with a final velocity dispersion of $\sim 8 \text{ km s}^{-1}$, much higher than Crater II, so there is no obvious conflict with our conclusions.

Satellite II, on the other hand, is less massive, and its halo has a less obvious core after being subjected to tides. This satellite also ‘tidally expands’ when subjected to extreme stripping (i.e. orbit V), leaving behind a stellar remnant with $M_* \sim 2.1 \times 10^4 M_\odot$, $\sigma_{\text{los}} \sim 2.5 \text{ km s}^{-1}$, and a large half-light radius, $r_h \sim 1.2 \text{ kpc}$ (see figs 4 and 9 in Frings et al. 2017). At face value, these properties are comparable to the properties of Crater II, and the large size,¹⁰ in particular, seems at odds with our results.

However, satellite II on orbit V is stripped so severely that the reported structural properties of the stellar component are unlikely to be robust. Evidence for this comes from apparent inconsistencies between the properties of the stellar and dark matter components. For example, the total circular velocity at the stellar half-light radius

refers to the location of an upturn in the surface brightness profile, whereas here we refer to an upturn in the σ_{los} profile.

¹⁰We note that ‘matching the half-light radius’ is no guarantee that the actual density profile will be matched. Including weakly bound particles would result in large estimates of r_h , but it would still fail to fit the density profile, unless the core radius is matched too. See discussion in Section 2.1.

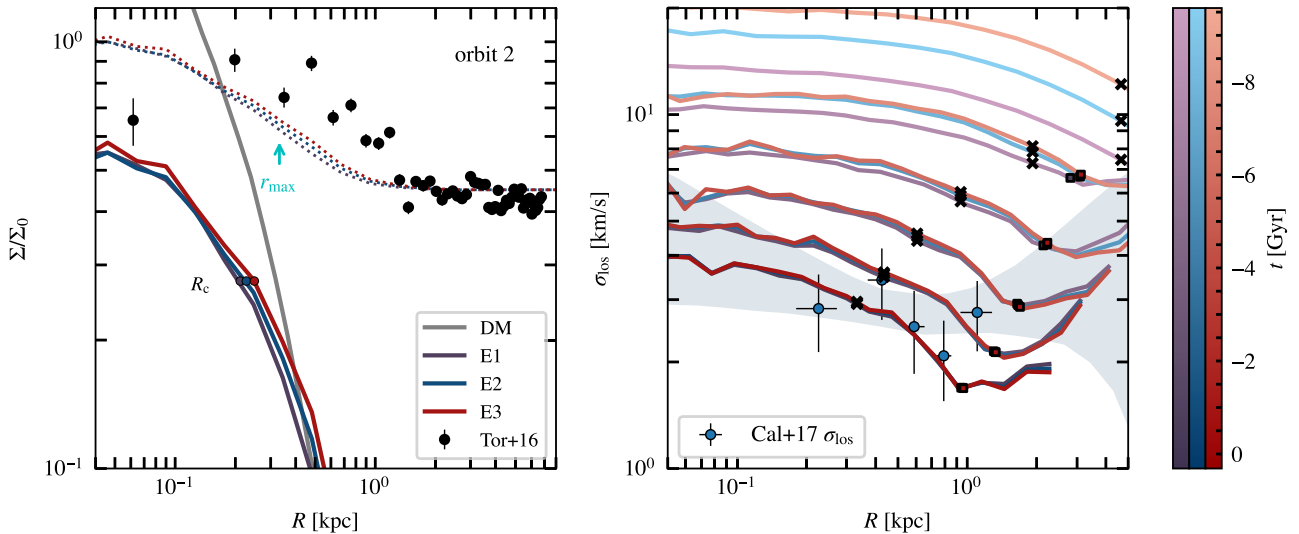


Figure 5. Left-hand panel: surface density profile of the bound remnant of stellar component models E1, E2, and E3, at the final snapshot identified for orbit 2, normalized to match the observed central density, Σ_0 . Note that the core radii are much smaller than observed, despite the wide range of initial core radii (~ 400 pc to ~ 2 kpc) sampled by the E1, E2, and E3 models. The characteristic radius of the remnant halo, r_{\max} , shown by an arrow, is a good indicator of the tidal truncation radius of the system. The dark matter halo projected density is drawn for reference in grey. Right-hand panel: line-of-sight velocity dispersion profiles for the bound constituents of the E1, E2, and E3 models, at each subsequent apocentric passage (see colour coding, where $t = 0$ is the ‘final’ snapshot of orbit 2). Model velocities roughly match the observed values at the final time (bottom profiles), although the sharp tidal truncation results in a sharp decline in σ_{los} outside r_{\max} (marked with an ‘X’ on each curve). The upturn in the outer regions is due to escaping stars moving away from the remnant. Squared symbols in each profile identify the location of that feature with the radius where the local crossing time equals the time elapsed since pericentre. See text for more details. Blue circles are observed values of σ_{los} for Crater II, taken from Caldwell et al. (2017).

is only $\sim 1 \text{ km s}^{-1}$ (see fig. 11 in Frings et al. 2017), over a factor of 4 smaller than the $V_c(r_h) \sim \sqrt{3} \sigma_{\text{los}} \approx 4.3 \text{ km s}^{-1}$ expected from simple mass estimators (see e.g. Wolf et al. 2010). A discrepancy this large is not expected with or without a core, and suggests that the reported properties of the stellar component are likely affected by either limited numerical resolution or substantial departures from equilibrium. Limited numerical resolution is certainly a possibility: A circular velocity of 1 km s^{-1} implies a total mass of $\sim 2.8 \times 10^5 M_\odot$ within r_h , which corresponds to fewer than ~ 200 dark matter particles.

Another possibility is that the estimated values of σ_{los} or r_h have been inflated by the inclusion of escaping or unrelaxed, weakly bound particles. If this were the case, then satellite II could not be compared with Crater II, whose inner regions are likely in dynamical equilibrium (see our earlier discussion). Unfortunately, there are not enough details in Frings et al. (2017) to fully track down the origin of the discrepancy with our results, but we stand by our conclusion that Crater II’s properties are inconsistent with those of a stellar component inhabiting the equilibrium tidal remnant of a cuspy NFW halo.

Limited resolution also helps to explain the findings of Applebaum et al. (2021), who report a number of simulated satellites with stellar mass ($\sim 10^5 M_\odot$), velocity dispersion ($\sim 3 \text{ km s}^{-1}$), and size (~ 1 kpc) comparable to Crater II, again apparently at odds with our findings. Cores cannot be invoked as an explanation here, since, according to these authors, cored halo profiles are only produced in their simulations in systems with $M_\star > 10^7 M_\odot$.

A more likely explanation is that the large sizes reported for Crater II-like satellites by these authors result because the progenitor haloes of Crater II-like systems are resolved with only $\sim 10^5$ dark matter particles, nearly 100 times fewer than in our runs. Indeed, limited resolution yields ‘tidal tracks’ that deviate systematically from the ‘EN21’ track shown in Fig. 4, resulting in artificially large values of r_{\max} in the case of heavy mass-loss.

We show this in Fig. 6, where we plot in grey the tidal track of a 10^5 -particle NFW halo otherwise identical to that used in our runs, placed on orbit 2. Solid black curves, as in Fig. 4, indicate the evolution of the 10^7 -particle halo. Circles along each track indicate the values of r_{\max} and V_{\max} at subsequent orbital apocentres. The 10^5 -particle halo follows a track that deviates from the EN21 track after the second pericentric passage, approaching a nearly constant value of $r_{\max} \approx 1$ kpc before (spuriously) fully disrupting before the fourth apocentre. Such deviation is well understood as a result of insufficient resolution. Indeed, the blue circles in Fig. 6 show the result of the evolution of many 10^5 -particle NFW haloes undergoing tidal stripping, taken from fig. A1 of Errani & Navarro (2021). These systems systematically deviate from the ‘EN21’ track to yield overestimated values of r_{\max} .

If, as discussed earlier, the size of ‘tidally limited’ galaxies is dictated by r_{\max} , then this would imply a typical size of about ~ 1 kpc for the remnants of $\sim 2 \times 10^9 M_\odot$ haloes such as those studied by Applebaum et al. (2021). The coincidence between this radius and the reported sizes of their Crater II-like candidates supports our view that the large reported sizes of those systems are unduly affected by limited numerical resolution.

Do our results then imply that Crater II’s halo had a core? One problem with this interpretation is that core formation in LCDM haloes usually results from vigorous inflows and outflows of gas during galaxy formation (Navarro et al. 1996a; Peñarrubia et al. 2012; Pontzen & Governato 2012; Benítez-Llambay et al. 2019), which are expected to occur only in galaxies more luminous than Crater II (Di Cintio et al. 2014). The model of Peñarrubia et al. (2012) suggests that, for the initial stellar and virial mass of Crater II adopted in this study, supernova feedback can produce only a very small core in the dark matter distribution of the order of $\sim 0.05 r_{\max}$.

Indeed, according to Tollet et al. (2016), sizeable cores are only expected in galaxies whose stellar mass-to-virial mass ratio exceeds a

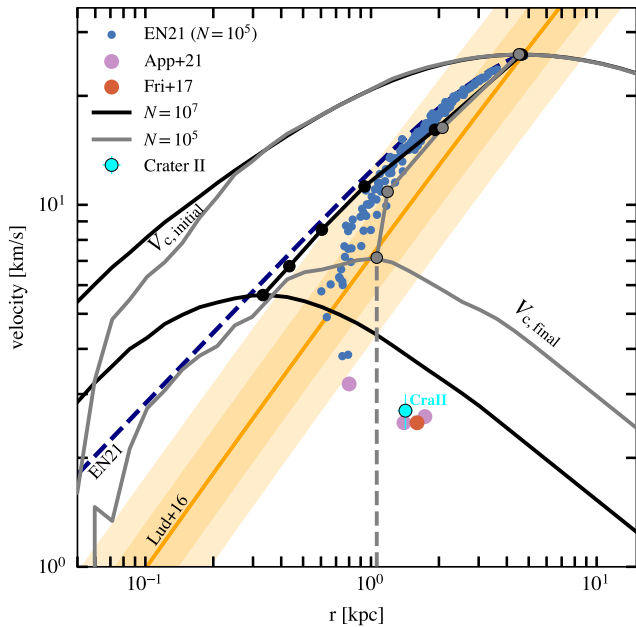


Figure 6. A comparison of tidal tracks of 10^5 - and 10^7 -particle realizations of the same NFW halo (solid grey and black curves, respectively). Circles along each track mark subsequent apocentric passages on orbit 2. Also shown are the quoted values of $r_{1/2}$ and σ_{los} from Applebaum et al. (2021) (purple) and Frings et al. (2017) (orange) for Crater II-like candidates, as well as those for Crater II (Caldwell et al. 2017; cyan symbol with error bar). Note that unlike the 10^7 -particle halo, the 10^5 -particle halo deviates from the EN21 tidal track and disrupts fully before its fourth apocentric passage. This deviation implies that lower resolution favours spuriously large values of r_{max} of the order of ~ 1 kpc, similar to those of the simulated Crater II-like candidates. This is confirmed by the results of Errani & Navarro (2021) for 10^5 -particle haloes (see their fig. A1), shown by the blue circles. See text for further details.

minimum value of $\sim 3 \times 10^{-4}$. This corresponds to a minimum stellar mass of $\sim 10^6 M_{\odot}$ for the progenitor halo considered here, roughly an order of magnitude higher than the stellar mass estimated for Crater II. This interpretation thus requires that the stellar component of Crater II was once much more massive than today, which seems unlikely given its low metallicity ($[\text{Fe}/\text{H}] = -2.10 \pm 0.08$; see e.g. fig. 8 in Ji et al. 2021).

Finally, Fattahi et al. (2018) used the Errani et al. (2015) tidal tracks to infer the properties of potential progenitors of Crater II, and concluded that Fornax-like systems stripped of more than 99 per cent of its stars could explain the unusual properties of Crater II. Those tracks, however, are applicable only for highly segregated stellar tracers and are not valid for systems as large and diffuse as Crater II. We present a detailed and updated discussion of such tracks in a separate contribution (Errani et al. 2021), which concludes, as we do here, that the large size of the Crater II dSph is inconsistent with that expected for an equilibrium tidal remnant of such a massive NFW halo.

3.4 Comparison with other dwarfs

The structural properties of Crater II are unusual, but not unique. We compare them to other dwarf galaxies in the Local Group in Fig. 7, where we plot the deprojected (3D) half-light radii, $r_{1/2}$, and circular velocities at that radius, $V_{1/2}$, for MW satellites (blue symbols), Andromeda satellites (cyan), as well as field dwarfs (brown). These were estimated from the observed projected half-light radius and line-of-sight velocity dispersion using the Wolf et al. (2010) mass

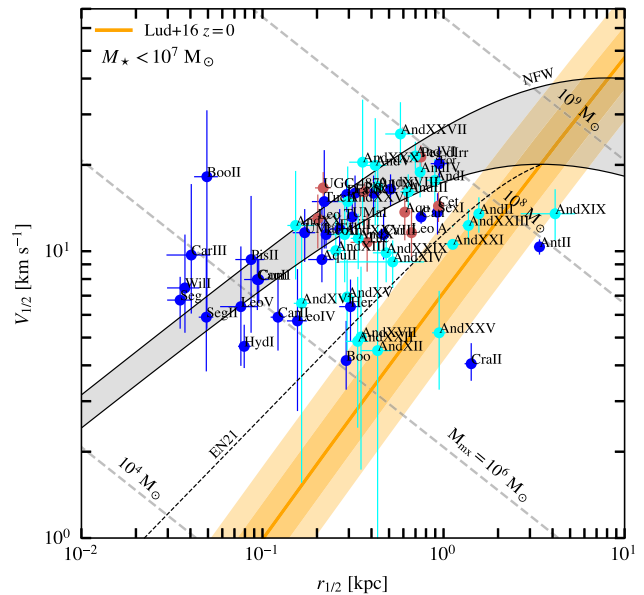


Figure 7. Satellites of the MW (dark blue) and of Andromeda (cyan), as well as Local Group field dwarfs (red) with $M_{\star} < 10^7 M_{\odot}$ (Fornax has also been added to the plot, for reference, although its stellar mass is slightly over the limit). 3D half-light radii, $r_{1/2}$, and circular velocities at that radius, $V_{1/2}$, are obtained using the mass estimator of Wolf et al. (2010) applied to the projected half-light radii and line-of-sight velocity dispersions of McConnachie (2012) (version 2021 January), updated with recent velocity dispersions for Antlia II, Crater II (Ji et al. 2021), Tucana (Taibi et al. 2020), And XIX (Collins et al. 2020), and And XXI (Collins et al. 2021). Systems with $M_{\star} < 10^7 M_{\odot}$ are expected to populate a narrow range of halo masses (Fattahi et al. 2018), shown schematically here by the grey band bracketing to NFW haloes with $V_{\text{max}} = 20$ and 40 km s^{-1} and average concentrations, taken from Ludlow et al. (2016, see the orange line with 1σ and 2σ scatter bands). Field dwarfs and satellites unaffected by tides should fall within the grey band. Tidally affected systems can drift below the grey band, but, as discussed in our work here, should stay to the left of the tidal track (labelled ‘EN21’) corresponding to that of the $V_{\text{max}} = 20 \text{ km s}^{-1}$ halo, the minimum mass of a halo likely to harbour a luminous dwarf (see e.g. Benitez-Llambay & Frenk 2020, and references therein).

estimator: $r_{1/2} = (4/3)R_{1/2}$ and $V_{1/2} \approx \sqrt{3}\sigma_{\text{los}}$. Data for Local Group galaxies come from the compilation of McConnachie (2012) (version 2021 January), with the addition of more recent data for some dwarfs, namely Tucana, Antlia 2, Crater II, And XIX, and And XXI (Collins et al. 2020, 2021; Taibi et al. 2020; Ji et al. 2021).

Dark matter-dominated dwarfs inhabiting unstripped, cuspy LCDM haloes with V_{max} in the range $(20, 40) \text{ km s}^{-1}$ (as expected for all dwarfs with $M_{\star} < 10^7 M_{\odot}$, regardless how faint) should lie roughly in the area highlighted in grey in Fig. 7 (Fattahi et al. 2018). Remarkably, the majority of Local Group dwarfs (~ 70 per cent) have error bars overlapping that region, which suggests that many of these galaxies indeed follow the simple expectations from LCDM simulations, with little evidence for tidal stripping playing a major role in their structure.

Systems above the grey area would suggest haloes more massive than $V_{\text{max}} = 40 \text{ km s}^{-1}$, but there are no dwarfs there, except for Bootes II, which has a remarkably high reported velocity dispersion for its size and luminosity (Koch et al. 2009). This has likely resulted from the inclusion of binary stars, which artificially inflated the estimate. Indeed, the analysis of Ji et al. (2016) concludes that Bootes II does not have a well-constrained velocity dispersion, and that earlier estimates are best regarded as upper limits.

Systems well below the grey area could in principle reflect haloes less massive than $V_{\max} = 20 \text{ km s}^{-1}$ (marked by the bottom boundary of the grey region), but this interpretation is disfavoured because such systems would be below the minimum halo mass required for hydrogen to cool efficiently (Efstathiou 1992; Gnedin 2000; Okamoto & Frenk 2009; Benitez-Llambay & Frenk 2020). The most likely interpretation of such systems (all of which are MW or M31 satellites) is that their dark matter content has been reduced by tides. Our discussion above, however, indicates that tidal effects cannot push systems in NFW-like haloes much below the dashed line labelled ‘EN21’, which traces the ‘tidal track’ of a $V_{\max} = 20 \text{ km s}^{-1}$ halo of average concentration.

The boundary delineated by this tidal track is a bit fuzzy because of the scatter in concentration around the mean value expected for LCDM (shown by the orange ‘error bands’ in Fig. 7), but it is clear that there exist a number of systems whose properties seem at odds with being tidal remnants of NFW-like haloes more massive than $V_{\max} \sim 20 \text{ km s}^{-1}$. Besides Crater II and Antlia II, for example, the M31 satellites And XXV and And XIX also have unusual sizes and velocity dispersions that are inconsistent with being tidal remnants of massive, cuspy LCDM haloes.

The interpretation of these outliers is not completely clear, but, if they are truly dark matter-dominated equilibrium systems, reconciling their properties with LCDM requires unpalatable choices. Going through all possible options in detail is beyond the scope of this contribution, but we note that the simplest option would be to just assume that these dSphs, despite their low stellar mass, had their dark matter cusp softened or removed by baryon effects. This could allow them to reach large sizes and low velocity dispersions, thus evading the boundary set by the ‘EN21’ tidal track shown in Fig. 7.

This solution, however, would raise the question of why many other dwarfs show no obvious sign of such core and lie comfortably within the grey region expected for dwarfs in LCDM. What singles out some of these dwarfs, in particular, to develop such cores? There are other alternatives, such as appealing to dark matter self-interactions to explain the presence of a ‘core’ rather than a ‘cusp’, but they all face a similar question: Why do only a few dwarfs (and not all, or even most) seem to require a ‘core’ to explain their low dark matter densities?

In this sense, the wide spread in dark matter content between dSphs shown in Fig. 7 evokes the puzzle presented by the ‘diversity’ in rotation curve shapes of dwarf irregular galaxies (dIrrs), rotationally dominated galaxies a couple of decades more luminous than the dSphs we consider here (Oman et al. 2015).

This diversity has not yet been fully explained, although it has elicited a number of proposed solutions, from baryon-induced modifications to the dark matter density profile (Navarro et al. 1996a; Pontzen & Governato 2012; Di Cintio et al. 2014), to modifications to the nature of cold dark matter, such as the inclusion of self-interactions (Rocha et al. 2013; Ren et al. 2019), to alterations of the laws of gravity, such as MOND (McGaugh 2016; Lelli et al. 2017), to, finally, the possibility that the data have been overinterpreted, and that the diversity is driven by underestimated uncertainties in the derivation of dIrr rotation curves (Oman et al. 2019). None of these proposed alternatives seems clearly favoured at present, as discussed by Santos-Santos et al. (2020).

4 SUMMARY AND CONCLUSIONS

We have used N -body simulations to study the effect of Galactic tides on the evolution of the Crater II dSph. The main aim of this work is to probe whether tidal effects may reconcile the unusual

size and kinematics of Crater II with that expected from LCDM cosmological hydrodynamical simulations. Our model assumes (i) that Crater II formed as a dark matter-dominated stellar system initially embedded in a cuspy NFW halo; (ii) that Crater II’s halo has initially a maximum circular velocity of $\sim 26 \text{ km s}^{-1}$ and average concentration, as expected from recent LCDM cosmological hydrodynamical simulations of the Local Group; and (iii) that Crater II stars may be modelled as a spherical system with no net rotation. Our main conclusions may be summarized as follows:

(i) The half-light radius of Crater II is poorly constrained due to the uncertain shape of the outer density profile. We therefore use the core radius when comparing simulation results against observations, which is less affected by assumptions about the outer profile shape.

(ii) The pericentric distance of Crater II on its orbit around the MW is not well constrained. We use the latest available data on the sky position, distance, radial velocity, and proper motions of Crater II, together with a Galactic potential model that matches the observed circular velocity at the solar radius, to find that the allowed pericentric distance lies roughly between ~ 10 and $\sim 50 \text{ kpc}$.

(iii) The negative Galactocentric radial velocity of Crater II, together with its large distance, implies that Crater II is just past its latest apocentric passage. More importantly, this also implies that the latest pericentric passage occurred roughly $\sim 1 \text{ Gyr}$ ago. This implies that the inner $\sim 1 \text{ kpc}$ of Crater II should be close to dynamical equilibrium, excluding models where its unusual properties are ascribed to substantial departures from equilibrium.

(iv) Assuming that Crater II initially formed in an NFW halo with virial mass $M_{200} = 2.7 \times 10^9 M_{\odot}$ (or $V_{\max} \approx 26 \text{ km s}^{-1}$, as suggested by results of the APOSTLE suite of LCDM Local Group simulations), we find that Galactic tides are able to reduce, in a Hubble time, the characteristic velocity of the remnant to $\sim 6 \text{ km s}^{-1}$ (as required to match Crater II’s low σ_{los}) if placed on an orbit with $r_{\text{peri}} \lesssim 15 \text{ kpc}$.

(v) The stellar components of the halo remnants tidally stripped as described above have core radii of the order of $\sim 200 \text{ pc}$ – almost four times smaller than observed. This is a general result for stellar components of NFW halo remnants, regardless of their initial radial extent. Crater II’s unusual size and velocity dispersion are thus inconsistent with the tidal evolution of a dwarf galaxy in cuspy NFW haloes.

(vi) Crater II’s unusual properties are shared by other satellite galaxies in the Local Group, like Antlia II, And XXV, and And XXI. None of these galaxies are consistent with a tidal interpretation in the standard LCDM scenario.

Reconciling Crater II-like systems with LCDM requires therefore some additional assumptions, none of them particularly appealing. The least unpalatable is that Crater II’s halo had a shallower cusp or inner density core formed during the assembly of the galaxy. It remains to be seen whether this is possible given the low stellar mass of the dwarf. Another possibility is that Crater II’s unusual properties arise because the system is quite far from equilibrium. This seems unlikely, given that Crater II is past its latest apocentric passage, and therefore its main body should have relaxed to equilibrium.

A further option is that the estimated photometric parameters of Crater II are somehow in error. This is also unlikely, given that the parameters have been estimated by independent groups (see e.g. the recent work of Moskowit & Walker 2020). Still, both these authors and the original Torrealba et al. (2016) discovery paper report a projected density profile that has an unexpected central *minimum* and a dynamic range of just a factor of ~ 2 in density. These limitations

urge further work to rule out the possibility that Crater II's core radius could be perhaps substantially smaller than reported so far.

Should none of these possibilities pan out, one would be forced to consider other alternatives, such as a halo of much lower mass and concentration than assumed in our work, or even the possibility that Crater II's unusual properties signal the need to consider modifying one or more of the foundational assumptions of the Λ CDM paradigm.

ACKNOWLEDGEMENTS

RE acknowledges support provided by a CITA National Fellowship and by funding from the European Research Council (ERC) under the European Unions Horizon 2020 research and innovation programme (grant agreement no. 834148). AF is supported by a UKRI Future Leaders Fellowship (grant no. MR/T042362/1). This work used the DiRAC@Durham facility managed by the Institute for Computational Cosmology on behalf of the STFC DiRAC HPC Facility (www.dirac.ac.uk). The equipment was funded by BEIS capital funding via STFC capital grants ST/K00042X/1, ST/P002293/1, ST/R002371/1, and ST/S002502/1, Durham University and STFC operations grant ST/R000832/1. DiRAC is part of the National e-Infrastructure.

DATA AVAILABILITY

The data underlying this paper will be shared on reasonable request to the corresponding author.

REFERENCES

- Aguilar L. A., White S. D. M., 1986, *ApJ*, 307, 97
 Amorisco N. C., 2019, *MNRAS*, 489, L22
 Applebaum E., Brooks A. M., Christensen C. R., Munshi F., Quinn T. R., Shen S., Tremmel M., 2021, *ApJ*, 906, 96
 Benítez-Llambay A., Frenk C., 2020, *MNRAS*, 498, 4887
 Benítez-Llambay A., Frenk C. S., Ludlow A. D., Navarro J. F., 2019, *MNRAS*, 488, 2387
 Borukhovetskaya A., Errani R., Navarro J. F., Fattahi A., Santos-Santos I., 2021, *MNRAS*, 509, 5330
 Bullock J. S., Johnston K. V., 2005, *ApJ*, 635, 931
 Caldwell N. et al., 2017, *ApJ*, 839, 20
 Collins M. L. M., Tollerud E. J., Rich R. M., Ibata R. A., Martin N. F., Chapman S. C., Gilbert K. M., Preston J., 2020, *MNRAS*, 491, 3496
 Collins M. L. M. et al., 2021, *MNRAS*, 505, 5686
 Di Cintio A., Brook C. B., Dutton A. A., Macciò A. V., Stinson G. S., Knebe A., 2014, *MNRAS*, 441, 2986
 Efsthathiou G., 1992, *MNRAS*, 256, 43P
 Einasto J., 1965, *Tr. Astrofiz. Inst. Alma-Ata*, 5, 87
 Errani R., Navarro J. F., 2021, *MNRAS*, 505, 18
 Errani R., Peñarrubia J., 2020, *MNRAS*, 491, 4591
 Errani R., Peñarrubia J., Tormen G., 2015, *MNRAS*, 449, L46
 Errani R., Navarro J. F., Ibata R., Peñarrubia J., 2022, *MNRAS*, 511, 6001
 Fattahi A. et al., 2016, *MNRAS*, 457, 844
 Fattahi A., Navarro J. F., Frenk C. S., Oman K. A., Sawala T., Schaller M., 2018, *MNRAS*, 476, 3816
 Frings J., Macciò A., Buck T., Penzo C., Dutton A., Blank M., Obreja A., 2017, *MNRAS*, 472, 3378
 Fritz T. K., Battaglia G., Pawlowski M. S., Kallivayalil N., van der Marel R., Sohn S. T., Brook C., Besla G., 2018, *A&A*, 619, A103
 Fu S. W., Simon J. D., Alarcón Jara A. G., 2019, *ApJ*, 883, 11
 Gnedin N. Y., 2000, *ApJ*, 542, 535
 Hayashi E., Navarro J. F., Taylor J. E., Stadel J., Quinn T., 2003, *ApJ*, 584, 541
 Hernquist L., 1990, *ApJ*, 356, 359
 Ji A. P., Frebel A., Simon J. D., Geha M., 2016, *ApJ*, 817, 41

- Ji A. P. et al., 2021, *ApJ*, 921, 32
 Kallivayalil N. et al., 2018, *ApJ*, 867, 19
 Koch A. et al., 2009, *ApJ*, 690, 453
 Kravtsov A., 2010, *Adv. Astron.*, 2010, 281913
 Lelli F., McGaugh S. S., Schombert J. M., Pawlowski M. S., 2017, *ApJ*, 836, 152
 Ludlow A. D., Bose S., Angulo R. E., Wang L., Hellwing W. A., Navarro J. F., Cole S., Frenk C. S., 2016, *MNRAS*, 460, 1214
 Macciò A. V., Frings J., Buck T., Penzo C., Dutton A. A., Blank M., Obreja A., 2017, *MNRAS*, 472, 2356
 McConnachie A. W., 2012, *AJ*, 144, 4
 McConnachie A. W., Venn K. A., 2020, *Res. Notes Am. Astron. Soc.*, 4, 229
 McGaugh S. S., 2016, *ApJ*, 832, L8
 McMillan P. J., 2011, *MNRAS*, 414, 2446
 Miyamoto M., Nagai R., 1975, *PASJ*, 27, 533
 Moskowit A. G., Walker M. G., 2020, *ApJ*, 892, 27
 Navarro J. F., 1990, *MNRAS*, 242, 311
 Navarro J. F., Eke V. R., Frenk C. S., 1996a, *MNRAS*, 283, L72
 Navarro J. F., Frenk C. S., White S. D. M., 1996b, *ApJ*, 462, 563
 Navarro J. F., Frenk C. S., White S. D. M., 1997, *ApJ*, 490, 493
 Okamoto T., Frenk C. S., 2009, *MNRAS*, 399, L174
 Oman K. A. et al., 2015, *MNRAS*, 452, 3650
 Oman K. A., Marasco A., Navarro J. F., Frenk C. S., Schaye J., Benítez-Llambay A., 2019, *MNRAS*, 482, 821
 Peñarrubia J., Navarro J. F., McConnachie A. W., 2008, *ApJ*, 673, 226
 Peñarrubia J., Navarro J. F., McConnachie A. W., Martin N. F., 2009, *ApJ*, 698, 222
 Peñarrubia J., Benson A. J., Walker M. G., Gilmore G., McConnachie A. W., Mayer L., 2010, *MNRAS*, 406, 1290
 Peñarrubia J., Pontzen A., Walker M. G., Koposov S. E., 2012, *ApJ*, 759, L42
 Plummer H. C., 1911, *MNRAS*, 71, 460
 Pontzen A., Governato F., 2012, *MNRAS*, 421, 3464
 Read J. I., Gilmore G., 2005, *MNRAS*, 356, 107
 Ren T., Kwa A., Kaplinghat M., Yu H.-B., 2019, *Phys. Rev. X*, 9, 031020
 Rocha M., Peter A. H. G., Bullock J. S., Kaplinghat M., Garrison-Kimmel S., Oñorbe J., Moustakas L. A., 2013, *MNRAS*, 430, 81
 Sanders J. L., Evans N. W., Dehnen W., 2018, *MNRAS*, 478, 3879
 Santos-Santos I. M. E. et al., 2020, *MNRAS*, 495, 58
 Santos-Santos I. M. E., Sales L. V., Fattahi A., Navarro J. F., 2021, preprint ([arXiv:2111.01158](https://arxiv.org/abs/2111.01158))
 Sawala T. et al., 2016, *MNRAS*, 457, 1931
 Springel V., 2005, *MNRAS*, 364, 1105
 Taibi S., Battaglia G., Rejkuba M., Leaman R., Kacharov N., Iorio G., Jablonka P., Zoccali M., 2020, *A&A*, 635, A152
 Tollet E. et al., 2016, *MNRAS*, 456, 3542
 Torrealba G., Koposov S. E., Belokurov V., Irwin M., 2016, *MNRAS*, 459, 2370
 Torrealba G. et al., 2019, *MNRAS*, 488, 2743
 van den Bosch F. C., Ogiya G., 2018, *MNRAS*, 475, 4066
 van den Bosch F. C., Ogiya G., Hahn O., Burkert A., 2018, *MNRAS*, 474, 3043
 van Dokkum P. G., Abraham R., Merritt A., Zhang J., Geha M., Conroy C., 2015, *ApJ*, 798, L45
 Walker M. G., Mateo M., Olszewski E. W., Peñarrubia J., Evans N. W., Gilmore G., 2009, *ApJ*, 704, 1274
 Wolf J., Martínez G. D., Bullock J. S., Kaplinghat M., Geha M., Muñoz R., Simon J. D., Avedo F. F., 2010, *MNRAS*, 406, 1220
 Woo J., Courteau S., Dekel A., 2008, *MNRAS*, 390, 1453

APPENDIX A: ORBITAL PATHS

This appendix presents the orbital paths of our Crater II halo on orbits 1, 2, and 3 as described in Section 2.3.

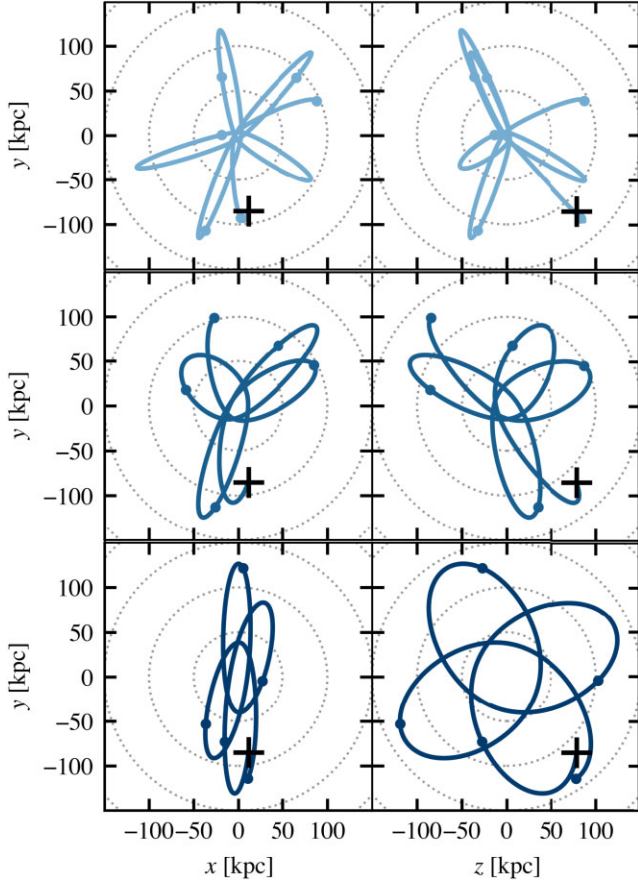


Figure A1. Projected trace on the X, Y (left-hand column) and Z, Y (right-hand column) planes of the three orbits considered in this study (see Table 2; the top row shows orbit 1, middle row – orbit 2, and bottom row – orbit 3). The current position of Crater II is indicated by a black cross. Intervals of 2 Gyr along the orbit are shown using filled circles. The Sun is at $(-8.3, 0, 0)$ in this coordinate system, with the velocity of the local standard of rest in the positive Y direction.

APPENDIX B: SUPPLEMENTARY TIDAL TRACKS

This appendix presents the supplementary tidal tracks referenced in Sections 3.2 and 3.3.

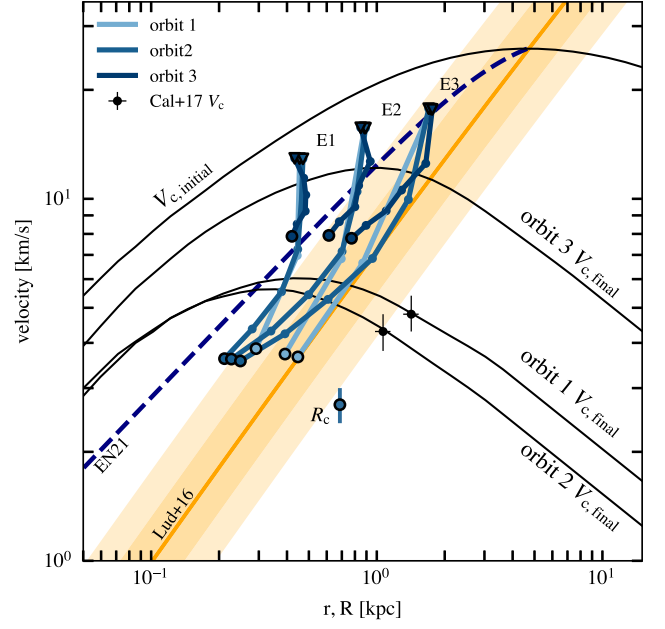


Figure B1. Like Fig. 4, but with the addition of stellar tracers (E1, E2, and E3 defined in Section 3.2) tracked through orbits 1 and 3 for comparison.

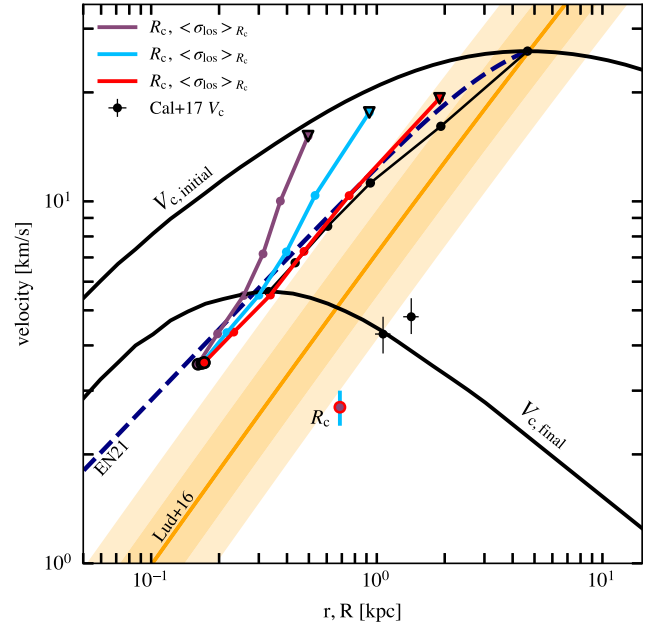


Figure B2. Like Fig. 4, but for stellar tracers which, at $t = 0$, have Einasto $\alpha = 0.5$ density profiles. The initial stellar core radii R_c are chosen to match those of the exponential tracers in Fig. 4. Also here, we observe that, in the limit of large tidal mass-losses, the stellar core radii and velocity dispersions evolve in sync with the characteristic size and velocity of the underlying DM halo.

This paper has been typeset from a \LaTeX file prepared by the author.