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# The LOFAR Two-metre Sky Survey Deep Fields

## The star-formation rate–radio luminosity relation at low frequencies

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### ABSTRACT

In this paper, we investigate the relationship between 150 MHz luminosity and the star-formation rate – the SFR- $L_{150\text{MHz}}$  relation – using 150 MHz measurements for a near-infrared selected sample of 118 517  $z < 1$  galaxies. New radio survey data offer compelling advantages over previous generation surveys for studying star formation in galaxies, including huge increases in sensitivity, survey speed, and resolution, while remaining impervious to extinction. The LOFAR Surveys Key Science Project is transforming our understanding of the low-frequency radio sky, with the 150 MHz data over the European Large Area Infrared Space Observatory Survey-North 1 field reaching an rms sensitivity of  $20\ \mu\text{Jy beam}^{-1}$  over  $10\ \text{deg}^2$  at 6 arcsec resolution. All of the galaxies studied have SFR and stellar mass estimates that were derived from energy balance spectral energy distribution fitting using redshifts and aperture-matched forced photometry from the LOFAR Two-metre Sky Survey (LoTSS) Deep Fields data release. The impact of active galactic nuclei (AGN) is minimised by leveraging the deep ancillary data in the LoTSS data release, alongside median-likelihood methods that we demonstrate are resistant to AGN contamination. We find a linear and non-evolving SFR- $L_{150\text{MHz}}$  relation, apparently consistent with expectations based on calorimetric arguments, down to the lowest SFRs  $< 0.01 M_{\odot}\ \text{yr}^{-1}$ . However, we also recover compelling evidence for stellar mass dependence in line with previous work on this topic, in the sense that higher mass galaxies have a larger 150 MHz luminosity at a given SFR, suggesting that the overall agreement with calorimetric arguments may be a coincidence. We conclude that, in the absence of AGN, 150 MHz observations can be used to measure accurate galaxy SFRs out to  $z = 1$  at least, but it is necessary to account for stellar mass in the estimation in order to obtain 150 MHz-derived SFRs accurate to better than 0.5 dex. Our best-fit relation is  $\log_{10}(L_{150\text{MHz}} / \text{W Hz}^{-1}) = (0.90 \pm 0.01) \log_{10}(\psi / M_{\odot}\ \text{yr}^{-1}) + (0.33 \pm 0.04) \log_{10}(M / 10^{10} M_{\odot}) + 22.22 \pm 0.02$ .

**Key words.** galaxies: star formation – radio continuum: galaxies

## 1. Introduction

Observations at radio wavelengths have great advantages for studying star formation across cosmic history, in particular the fact that they are impervious to the effects of the dust obscuration that blights star-formation rate (SFR) measures at optical wavelengths (e.g. Kennicutt 1998). In addition, as we embark upon the construction of the Square Kilometre Array (SKA; e.g. Carilli & Rawlings 2004; Dewdney et al. 2009), radio observations are undergoing an explosion of capabilities, including huge

increases in survey speed, spatial resolution, and sensitivity. The Low Frequency Array (LOFAR; van Haarlem et al. 2013) is already revolutionising our understanding of the low-frequency radio sky, and the LOFAR Surveys Key Science Project (LSKSP; Röttgering et al. 2011) aims to survey the entire northern sky with an unprecedented combination of sensitivity and angular resolution. Huge progress has already been made: the first data release (DR1) of the LOFAR Two-metre Sky Survey (LoTSS: Shimwell et al. 2017, 2019; Duncan et al. 2019; Williams et al. 2019) covered an area of  $424\ \text{deg}^2$  with a median sensitivity of

71  $\mu\text{Jy}$  at 150 MHz and with 6 arcsec resolution, while the forthcoming second data release will cover 5200  $\text{deg}^2$  of the northern sky with similar sensitivity. Within the LSKSP, the Low Band Antenna (LBA) is also being used to carry out a sister survey to LoTSS at 60 MHz: the LOFAR LBA Sky Survey (LoLSS; De Gasperin et al. 2021).

In addition to the 150 MHz survey that covers the entire northern sky, the LSKSP also produces deeper observations of multiple 10  $\text{deg}^2$ -scale regions with the best multi-wavelength data, known as the LoTSS Deep Fields. The first three of these fields, Boötes, European Large Area Infrared Space Observatory Survey-North 1 (ELAIS-N1), and Lockman Hole (which are the subject of the first deep field data release and are fully described in four papers: Tasse et al. 2021; Sabater et al. 2021; Kondapally et al. 2021; and Duncan et al. 2021), reach sensitivities between 20 and 35  $\mu\text{Jy}$  rms, a factor of two to three times deeper than the standard LoTSS observations. At these flux densities, the source counts are increasingly dominated by star-forming galaxies (e.g. Retana-Montenegro et al. 2018; Williams et al. 2019) in which the synchrotron radiation is attributed to electrons accelerated by the remnants of supernovae that are the end points of the evolution of short-lived massive stars. The association between particle acceleration and the supernova rate has been used to directly calibrate the synchrotron luminosity as an SFR indicator (e.g. Condon 1992; Cram et al. 1998).

Further support for the use of radio frequency observations to study star formation in galaxies comes in the form of the far-infrared to radio correlation (FIRC), which many works have shown to be a tight and constant relationship that persists over several orders of magnitude in luminosity (e.g. van der Kruit 1971; de Jong et al. 1985; Helou et al. 1985; Yun et al. 2001; Appleton et al. 2004; Jarvis et al. 2010; Ivison et al. 2010a,b; Sargent et al. 2010; Bourne et al. 2011; Delhaize et al. 2017). However, relying on the FIRC to underpin the SFR calibration of radio luminosity is sub-optimal, in the sense that although a constant FIRC can be explained by calorimetry arguments, conspiracies (i.e. the precise balance between disparate phenomena such as the cosmic ray electron escape fraction and the optical depth to UV photons; e.g. Lisenfeld et al. 1996; Bell 2003; Lacki et al. 2010) are required to explain the correlation itself. Furthermore, there is good empirical evidence that the FIRC is not constant, varying in different galaxy types (e.g. Molnár et al. 2018; Read et al. 2018) as a function of dust temperature (Smith et al. 2014) and redshift (e.g. Magnelli et al. 2015).

Clearly then, it is essential that we use the best observations available to test the efficacy of radio continuum data as an SFR indicator and determine the best functional form to use. Many observational works have looked directly at the SFR – radio luminosity relation (e.g. Condon 1992; Cram et al. 1998; Garn et al. 2009; Kennicutt et al. 2009; Murphy et al. 2011; Tabatabaei et al. 2017), finding broad support for proportionality. On the other hand, Bell (2003) determined a downturn in the relation at low SFRs (in the sense of lower luminosity per unit SFR) on the pragmatic basis of a decreasing non-thermal fraction at 1.4 GHz in less luminous sources; this was couched within a wider discussion of possible physical reasons for the variation, including the possibility of increasing cosmic ray escape in this regime (e.g. Chi & Wolfendale 1990; Murphy et al. 2008). Similarly, Lacki & Thompson (2010) and Murphy (2009) predict a deviation in the relation at high- $z$  due to inverse Compton losses that are the result of a significantly increased cosmic microwave background photon density at earlier points in cosmic history. Davies et al. (2017) used data from the Galaxy And Mass Assembly (GAMA; Driver et al. 2011) and Faint Images

of the Radio Sky at Twenty Centimeters (FIRST; Becker et al. 1995) surveys to study the SFR – radio luminosity relation, while Hodge et al. (2008) used the reprocessed Sloan Digital Sky Survey (SDSS) spectroscopy (York et al. 2000) in the Max Planck Institute for Astrophysics and John Hopkins University catalogue (MPA-JHU; Brinchmann et al. 2004) alongside FIRST data to do likewise; both studies uncovered clear evidence for a super-linear slope<sup>1</sup>. Each of these studies present strong evidence for deviation from the linear (i.e. gradient of unity) form that would be expected on the basis of calorimetric models.

In addition to the survey speed, sensitivity, and resolution benefits that LOFAR provides, LoTSS observations have far superior sensitivity to extended emission compared to FIRST due to the short baselines sampled (e.g. Hardcastle et al. 2019a; Mahatma et al. 2019; Tasse et al. 2021). Furthermore, the low frequencies may also be preferable for studying the relationship between the SFR and radio luminosity since free-free emission (which can dominate at gigahertz frequencies, e.g. Condon 1992; Murphy et al. 2011) makes a negligible contribution to the 150 MHz luminosity. Low-frequency observations are therefore potentially “cleaner” than the much better studied gigahertz frequency range (provided that we do not observe to such low frequencies that the spectra become self-absorbed, e.g. Kellermann & Owen 1988). With the advent of LOFAR, several works have looked at this issue at low frequencies (e.g. Brown et al. 2017; Calistro Rivera et al. 2017; Gürkan et al. 2018; Wang et al. 2019). Brown et al. (2017) found a slightly super-linear relationship between the SFR and the 150 MHz luminosity (hereafter the “SFR- $L_{150\text{MHz}}$ ” relation) with a gradient of  $1.14 \pm 0.05$ , using 150 MHz data from the Tata Institute of Fundamental Research Giant Metrewave Radio Telescope Sky Survey (TGSS; Intema et al. 2017). Wang et al. (2019) found an even steeper relationship (gradient of 1.3–1.4) using a 60  $\mu\text{m}$ -selected sample from the revised Infrared Astronomical Satellite (IRAS) Faint Source Survey Redshift Catalogue (Wang et al. 2014) that they matched to LOFAR data from LoTSS DR1. Calistro Rivera et al. (2017) used a 150 MHz-selected sample of 750  $z < 2.5$  star-forming galaxies identified in LoTSS data over 7  $\text{deg}^2$  of the Boötes field, alongside SFRs from spectral energy distribution (SED) fitting, to find evidence for significant evolution in the FIRC, curved radio spectra, and an increase in the observed radio luminosity for a given SFR at higher redshift.

Of particular relevance, Gürkan et al. (2018, hereafter G18) studied SFR- $L_{150\text{MHz}}$  using a spectroscopically classified sample from the MPA-JHU catalogue containing  $\sim 15\text{k}$  galaxies with the first LOFAR observations of the *Herschel* Astrophysical Terahertz Large Area Survey (H-ATLAS) North Galactic Pole (NGP) field (from Hardcastle et al. 2016, reaching an rms of 100  $\mu\text{Jy}$  near the centre of the field but with non-uniform coverage). G18 used SFRs and stellar masses based on 14-band SED fitting with the MAGPHYS code (da Cunha et al. 2008) and found compelling evidence for (a) significant scatter about the maximum-likelihood SFR- $L_{150\text{MHz}}$  relation, (b) a strong preference for a mass dependence in SFR- $L_{150\text{MHz}}$ , and, perhaps most intriguingly (c) an upturn towards larger radio luminosity at SFRs  $\log_{10}(\psi/M_{\odot}\text{yr}^{-1}) < 0$ , in the direction opposite to that expected from calorimetry arguments. This  $L_{150\text{MHz}}$  excess was also found by Read et al. (in prep.) using 150 MHz data from LoTSS DR1 over the *Hobby-Eberly* Telescope Dark Energy Experiment (HETDEX) field alongside SFRs derived from the

<sup>1</sup> We note that the slope of  $0.75 \pm 0.03$  quoted by Davies et al. (2017) uses the inverse definition of the SFR – radio luminosity relation, which is why the quoted value is lower than unity.

MPA-JHU catalogue. The cause of this extra radio luminosity in the galaxies with the lowest SFRs is of potentially great interest since the rapid increase in radio survey capabilities means that future radio surveys will access this regime ever more readily. Possible mechanisms for providing additional cosmic rays at low SFRs include pulsars, type Ia supernovae, residual contamination by active galactic nuclei (AGN), and varying magnetic field properties (see also [Sudoh et al. 2020](#)).

Putting these issues to one side, perhaps the principal reason against using low-frequency radio observations for studying star formation is that other physical phenomena can also be bright at radio frequencies. The main issue for this type of study is activity due to AGN, where the synchrotron radio luminosity is linked not to star formation but to accretion, with the necessary relativistic electrons instead coming from particle acceleration in jets. While [Sabater et al. \(2019\)](#) has shown that the most massive radio-selected galaxies are always radio loud, AGN are far from ubiquitous in optically selected samples. However, AGN contamination is virtually impossible to eradicate completely from samples of star-forming galaxies at any wavelength due to the fact that AGN can be highly dust obscured (e.g. [Antonucci 1993](#); [Lacy et al. 2004](#); [Martínez-Sansigre et al. 2005](#)) since accretion processes are variable (e.g. [Read et al. 2020](#)) and since AGN are highly multi-modal (e.g. [Hardcastle et al. 2007](#); [Best & Heckman 2012](#)). In addition, radio AGN have a very large range of power (e.g. [Hardcastle et al. 2019b](#)), extending down to values that would be typical of star-forming galaxies (e.g. [Sadler et al. 2002](#); [Hardcastle et al. 2016](#); [Lofthouse et al. 2018](#)), which makes simple cuts in luminosity suboptimal.

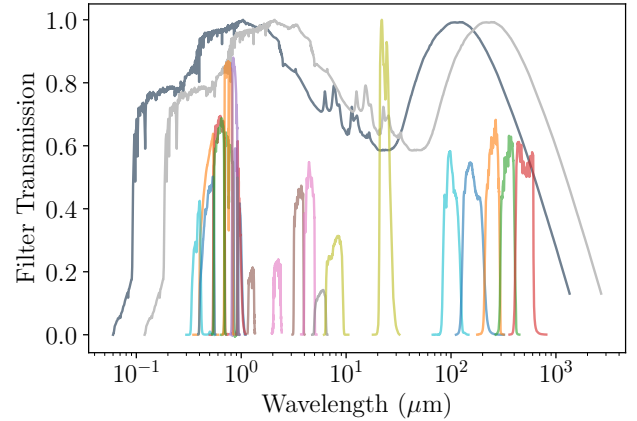
Despite the issues that AGN contamination presents, the potential gains from using radio survey data for studying star formation are significant, and this is reflected by the wide range of authors who have done so to date (e.g. [Haarsma et al. 2000](#); [Hopkins et al. 2003](#); [Pannella et al. 2009, 2015](#); [Karim et al. 2011](#); [Zwart et al. 2014](#); [Bonato et al. 2017](#); [Novak et al. 2017](#); [Upjohn et al. 2019](#); [Leslie et al. 2020](#)). In this paper, we revisit the low-frequency radio luminosity SFR relation using data from the LOFAR deep fields. The structure of this work is as follows: In Sect. 2, we introduce the different data sets that we use for our analysis, while in Sect. 3 we discuss our analysis and present the results. In Sect. 4, we present some concluding remarks. We assume a standard cosmology with  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_m = 0.3$ , and  $\Omega_\Lambda = 0.7$ .

## 2. Data

In this work, we focus on the  $6.7 \text{ deg}^2$  of the ELAIS-N1 field, where the best multi-wavelength data exist ([Kondapally et al. 2021](#)). We describe here the key data sets that we use in Sects. 2.1–2.5.

### 2.1. Multi-wavelength data

We used the aperture-matched photometry from [Kondapally et al. \(2021\)](#). In the ELAIS-N1 field, the photometric bands include the *Spitzer* Adaptation of the Red-sequence Cluster Survey (SpARCS) *u*-band ([Wilson et al. 2009](#); [Muzzin et al. 2009](#)), Panoramic Survey Telescope and Rapid Response System (PanSTARRS) grizy ([Chambers et al. 2016](#)), and NB921 narrow-band data from Hyper Suprime-Cam Subaru Strategic Program (HSC-SSP) public data release 1 ([Aihara et al. 2018](#)), *J*- and *K*-band data from the Deep Extragalactic Survey (DXS; [Lawrence et al. 2007](#)), Infrared Array Camera (IRAC) channel 1-4 data from the *Spitzer* Wide-Area Infrared



**Fig. 1.** Filter curves of the assembled multi-wavelength data set (coloured lines), overlaid with a galaxy template from [Smith et al. \(2012\)](#). The SED has been normalised to have a peak flux of 1, and it is shown both at  $z=0$  (in grey) and with the wavelength axis shifted to  $z=1$  (in silver) to give the reader an impression of how the excellent multi-wavelength coverage in ELAIS-N1 samples a galaxy’s SED.

Extragalactic Survey (SWIRE; [Lonsdale et al. 2003](#)), and channel 1 and 2 data from the *Spitzer* Extragalactic Representative Volume Survey (SERVS; [Mauduit et al. 2012](#)). These data have been reprocessed onto a common pixel scale and used in combination with state-of-the-art photometric redshift information from [Duncan et al. \(2021\)](#) to produce consistent aperture forced photometry for every source in the catalogue. In this paper, we use the compiled spectroscopic redshifts where available and reliable, and the [Duncan et al. \(2021\)](#) photometric redshift estimates otherwise (1.5% of our sample have spectroscopic redshifts). We have also used the SWIRE  $24 \mu\text{m}$  data as provided by the *Herschel* Extragalactic Legacy Project (HELP; [Shirley et al. 2019](#)); this also includes far-infrared data in the 100, 160, 250, 350, and  $500 \mu\text{m}$  bands from the *Herschel* Multi-tiered Extragalactic Survey (HerMES; [Oliver et al. 2012](#)) that we probabilistically de-blended using the XID+ tool ([Hurley et al. 2017](#)) as described by [McCheyne et al. \(in prep.\)](#). To demonstrate the full wavelength coverage of the extensive multi-wavelength data set that has been assembled, we show in Fig. 1 the filter curves overlaid on an indicative galaxy spectrum template from [Smith et al. \(2012\)](#). The template is shown both at  $z=0$  and with the wavelength axis shifted to  $z=1$  to illustrate the range spanned by our sample; it is clear that there is excellent coverage all the way from the near-ultraviolet to the far-infrared wavelengths.

### 2.2. Sample definition

To avoid the possible influence of radio selection on our results (see Appendix B for further details), we begin with a sample identified in the  $3.6 \mu\text{m}$  data from SWIRE. These data are not only very sensitive, but they are also less susceptible to the influence of the dust obscuration that could introduce bias into samples identified at shorter wavelengths. We selected those sources with measured  $3.6 \mu\text{m}$  flux density  $> 10 \mu\text{Jy}$ , approximately equal to the  $5\sigma$  detection threshold in the SWIRE  $3.6 \mu\text{m}$  data and as recommended on the SWIRE webpages<sup>2</sup>. We included only those sources with the necessary flags in the band-merged catalogue, as recommended in [Kondapally et al. \(2021\)](#), to ensure we used only those sources with the highest-quality photometry,

<sup>2</sup> <https://irsa.ipac.caltech.edu/data/SPITZER/SWIRE/>

the best wavelength coverage, and the most reliable photometric redshifts<sup>3</sup>. This leaves us with a parent sample of 142 037  $3.6\ \mu\text{m}$ -selected sources at  $z < 1$ , where [Duncan et al. \(2021\)](#) show that the photometric redshifts have an average scatter of  $\sigma_{\text{NMAD}} < 0.04(1+z)$  and an outlier fraction around 5%<sup>4</sup>.

### 2.3. LOFAR data

We used the new deep LOFAR observations of the ELAIS-N1 field ([Sabater et al. 2021](#)). They cover an area of  $10\ \text{deg}^2$  that has an rms below  $30\ \mu\text{Jy}$ , reaching  $20\ \mu\text{Jy}$  in the deepest regions, which makes them the most sensitive data in the LoTSS Deep Fields DR1. Although the area covered is significantly narrower than that studied in G18, the 150 MHz data are around five times deeper. Importantly, due to the very high quality of the multi-wavelength data in this field, more than 97 percent of the sources in the ELAIS-N1 LoTSS 150 MHz catalogue have counterparts in at least one band, which have been reliably identified using a new colour-dependent implementation ([Kondapally et al. 2021](#)) of the likelihood ratio method ([Sutherland & Saunders 1992](#); [Smith et al. 2011](#); [McAlpine et al. 2012](#); [Nisbet 2018](#); [Williams et al. 2019](#)). We used the [Sabater et al. \(2021\)](#) catalogue total 150 MHz flux densities and uncertainties where they exist; however, for the  $\sim 91$  percent of IRAC sources that are not identified as the counterparts of sources in the 150 MHz catalogue, we used pixel flux densities, specifically the flux density measured in the 150 MHz map at the pixel corresponding to the coordinates of each IRAC source. These values indicate the maximum likelihood estimate of the integrated flux density for sources that are unresolved on the scale of the 6 arcsec LoTSS beam, and they are strongly preferred over aperture photometry in these deep data since they are less susceptible to influence from neighbouring sources. Including these sources is especially important since, although they are not individually detected (in the sense that they have a signal-to-noise ratio ( $S/N$ )  $< 3$ ), they are numerically dominant and together can provide significant diagnostic power, as we shall demonstrate. To limit the potential influence of IRAC sources with 150 MHz emission more extended than 6 arcsec, we only considered the 130 689 sources at  $z > 0.05$ . The [Kondapally et al.](#) catalogue includes size information for sources in the  $\chi^2$  image used for source detection, which we have used to estimate the extent of the sources in our sample. The mean full width at half maximum (FWHM) of galaxies along their major axes decreases from 2.7 arcsec at the lowest redshifts to around 1.2 arcsec by  $z = 1$ . In practice, therefore, this cut does not have a significant impact on our results since  $>99.5\%$  of  $z < 1$  sources – and  $>98\%$  even at  $z < 0.2$  – have an FWHM smaller than 6 arcsec in the multi-wavelength data. We also include 150 MHz flux density uncertainties measured from the corresponding pixel of the rms map.

### 2.4. SFR and mass estimation

We used the panchromatic energy balance SED fitting code MAGPHYS to fit the multi-wavelength matched-aperture photometry and determine the physical properties of each source. MAGPHYS is fully described in [da Cunha et al. \(2008\)](#), but to summarize: it uses an energy balance criterion to link the stellar emission that dominates at optical and near-infrared (ONIR)

**Table 1.** Percentage of the area covered in this study by each set of ancillary data used for the SED fitting, along with the mean percentage of sources detected at  $\geq 3\sigma$  in the most sensitive band of each data set.

Band	% Coverage	% $\geq 3\sigma$
SpARCS u	100	68
PanSTARRS grizy	100	100
HSC grizy, nb921	94	94
DXS JK	100	100
SWIRE IRAC ch1	100	100
SWIRE IRAC ch 2-4	100	96
SERVS IRAC 1 & 2	34	34
MIPS $24\ \mu\text{m}$	85	54
<i>Herschel</i> PACS	100	1
<i>Herschel</i> SPIRE	100	15

**Notes.** Information regarding the depth of the data in each band is available in [Kondapally et al. \(2021\)](#).

wavelengths with the dust emission that dominates in the far-infrared. The stellar emission is modelled using the templates from [Bruzual & Charlot \(2003\)](#) and assuming an initial mass function (IMF) from [Chabrier \(2003\)](#), attenuated using a two-component dust model from [Charlot & Fall \(2000\)](#); the energy absorbed is then re-radiated using a multi-component dust model (including dust grains of various sizes and temperatures, as well as polycyclic aromatic hydrocarbons). The energy balance criterion is used to combine an ONIR library based on 50 000 SEDs, which represent exponentially declining star-formation histories with stochastic bursts superposed, with a library of 50 000 dust SEDs with a realistic range of physical properties. The energy balance criterion is designed to ensure that only physical combinations of the stellar and dust SEDs are used to model the input photometry. MAGPHYS does not account for possible additional dust heating due to AGN (though see e.g. [Berta et al. 2013](#), for an attempt to include it). As well as producing best-fit SEDs for every source, we also obtained best-fit physical parameters with marginalised probability distribution functions (PDFs) for each parameter. These PDFs were used to derive median-likelihood parameter estimates (which are the values corresponding to the 50th percentile of the PDF) as well as the 16th and 84th percentiles, which are equivalent to the  $\pm 1\sigma$  values for each parameter in the limit of Gaussian statistics.

Since the multi-wavelength data over the ELAIS-N1 field have been compiled from a wide range of sources with different coverage (see Sect. 2.1), not every source has measured photometry in every band. Table 1 shows the fraction of sources in our sample that have coverage and/or a  $\geq 3\sigma$  detection in each band, while Table 2 shows the fraction of sources that have measured photometry and  $\geq 3\sigma$  detections in at least  $N$  photometric bands. While Table 1 highlights that 100% of our sample have *Herschel* coverage, only  $\sim 15\%$  are detected at  $\geq 3\sigma$  in the most-sensitive  $250\ \mu\text{m}$  band, with  $22\% \geq 2\sigma$  and  $43\% \geq 1\sigma$ . Indeed, only  $\sim 65\%$  of the sample have been assigned *Herschel* photometry by XID+; the fact that the remaining  $3.6\ \mu\text{m}$  sources were not assigned any measurable *Herschel* flux density, despite the absence of an SED prior in the version of XID+ we used (cf. [Pearson et al. 2017](#)), is a strong indication that they are fainter than the confusion noise in each of the Photodetector Array Camera and Spectrometer (PACS) and SPIRE bands ([Hurley et al. 2017](#)). To include this “upper limit” information in the MAGPHYS parameter estimation, we assigned these sources uncertainties in the PACS and

<sup>3</sup> As noted in [Kondapally et al. \(2021\)](#), in the ELAIS-N1 field, the necessary flags to apply are `flag_clean=1` & `flag_overlap=7`.

<sup>4</sup> The scatter is measured in terms of the normalised median absolute deviation, which is defined as  $\sigma_{\text{NMAD}} = 1.48 \times \text{median}(|\Delta z|/(1+z_{\text{spec}}))$ , while outliers are defined as those sources where  $|\Delta z|/(1+z_{\text{spec}}) > 0.15$ .

**Table 2.** Percentage of sources in our parent sample with measured photometry (middle column) and a  $\geq 3\sigma$  detection in at least  $N$  bands.

Number of bands	% with photometry	% $\geq 3\sigma$
7	100	100
8	100	99
9	100	98
10	100	96
11	100	94
12	99	92
13	98	89
14	98	86
15	98	77
16	98	62
17	98	44
18	84	28
19	72	16
20	67	8
21	61	4
22	61	2
23	61	1
24	25	0
25	21	0

**Notes.** For example, 98% of our sample has measured photometry in 17 or more bands, and more than half have  $\geq 3\sigma$  detections in at least 16 bands.

SPIRE bands equal to the median uncertainty for the sources that do have measured flux densities (these values are 10.3, 14.1, 2.6, 2.8, and 3.7 mJy in the 100, 160, 250, 350, and 500  $\mu\text{m}$  bands, respectively), alongside a small flux density (equal to 0.1% of the median uncertainty)<sup>5</sup>.

The principal quantities of interest for this work are the stellar mass and the SFR averaged over the last 100 Myr, for which MAGPHYS estimates have been shown to be reliable in a range of different situations (e.g. Hayward & Smith 2015; Smith & Hayward 2018; Dudzevičiūtė et al. 2020). To ensure that the MAGPHYS parameters are as reliable as possible, we further refined our sample to include only those 123 425 galaxies for which MAGPHYS has been able to produce an acceptable fit. We define “acceptable” to mean that the  $\chi^2$  parameter that compares the goodness of the fit between the model and the observed photometry is below the 99 percent confidence threshold derived for the number of bands of photometry available for each object, as discussed in Smith et al. (2012). Though we are unable to repeat their emission line classification, our sample is roughly a factor of eight larger than the one used by G18. In addition, the multi-wavelength data set is far better; the HSC  $i$ -band data in ELAIS-N1 are four magnitudes deeper (and the comparison is similar in the other HSC bands), while the IRAC 3.6  $\mu\text{m}$  data are  $>20$  times as sensitive as the closest comparable band in that work. Not only are the individual observations more sensitive, but there are also many more bands of photometry available for us to use. We demonstrate this in Table 2, which shows that more than half of our sample has photometry in 23 or more bands, as compared to the maximum of 14 that were available in the H-ATLAS NGP area used by G18.

<sup>5</sup> We repeated all of the following analysis both with and without including these sources in the analysis, and the results do not show significant variation.

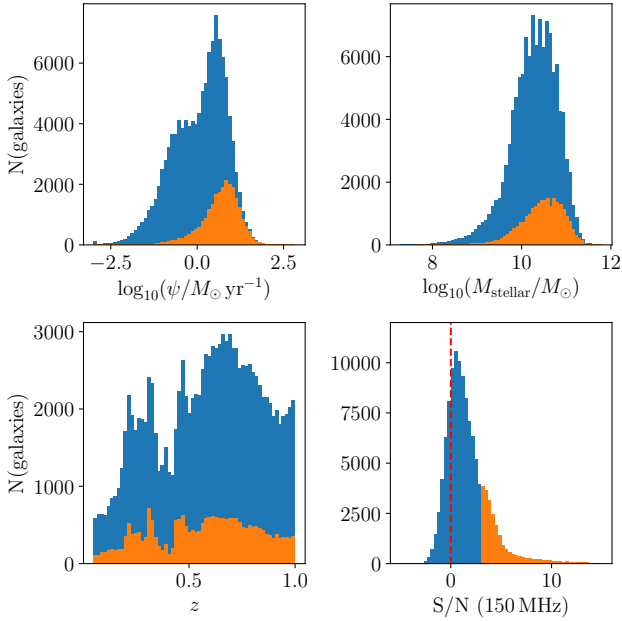
## 2.5. Sample properties and AGN contamination

As recommended by Duncan et al. (2021), we removed: (a) the most obvious AGN in our mass-selected sample using the flags supplied in the input catalogue, which include sources in the Million Quasar Catalog (Flesch 2019) and sources spectroscopically classified as AGN in the literature, (b) those sources that fall into the Donley et al. (2012) infrared colour space dominated by AGN, and (c) those with bright X-ray counterparts (see Duncan et al. 2021 for details); this left us with 122 646 sources. However, for those galaxies with 150 MHz detections, we also removed additional AGN that were identified using the method from Best et al. (in prep.), which relies on the results of a comprehensive multi-wavelength SED fitting analysis, using our MAGPHYS results alongside those derived using BAGPIPES (Carnall et al. 2018), AGNFITTER (Calistro Rivera et al. 2016), and CIGALE (Burgarella et al. 2005). The details of the Best et al. method are complex, but the general idea is to compare the results of those SED fitting codes that account for AGN (AGNFITTER and CIGALE) with those that are focussed on normal star-forming and passive galaxies (MAGPHYS and BAGPIPES) considering all of the available information including comparing the goodness-of-fit that each code produces, accounting for the best-fit AGN fractions, and identifying sources with a clear radio excess. A comparison between the Best et al. (in prep.) AGN flagging procedure and the results of removing the bad MAGPHYS fits using the aforementioned  $\chi^2$  threshold method from Smith et al. (2012) suggests that applying the MAGPHYS threshold alone removes  $\sim 96$  percent of the flagged AGN. This is especially useful for this work since the Best et al. (in prep.) AGN flags have only been derived for the sources detected in the 150 MHz catalogue<sup>6</sup>. After applying these cuts, we were left with 120 232  $z < 1$  galaxies, of which 9298 (25 777) are detected at  $\geq 5\sigma$  ( $\geq 3\sigma$ ) in the deep LoTSS data. We also identified 1715 further sources that were not flagged as AGN in the Best et al. radio excess sample but which had a clear radio excess in our pixel flux density measurements ( $>5\sigma$  150 MHz detections and  $L_{150\text{MHz}}$  more than 1 dex larger than the SFR- $L_{150\text{MHz}}$  relation from G18; this is 1.4 percent of our sample), which indicates they are likely undiagnosed AGN; we thus removed them from our sample, leaving us with 118 517 galaxies on which to base our analyses. We repeated the analyses and note that the results did not significantly change regardless of whether or not we included these sources, whether or not we used a 1 dex excess or a 0.7 dex radio excess to identify residual AGN, or whether or not we instead used the mass-dependent SFR- $L_{150\text{MHz}}$  relation from G18.

Histograms showing the median-likelihood SFR, stellar mass, redshift, and 150 MHz S/N distribution for the galaxies in our sample are shown in Fig. 2. In Fig. 3, we display the stellar mass (top) and SFR (bottom) distributions as a function of redshift, colour-coded by the number of galaxies in each bin.

In Fig. 4, we show the location of the SFR and 150 MHz luminosity for those 25 777 sources that are detected at  $>3\sigma$  in the deep LoTSS data, overlaid with the mass-independent SFR- $L_{150\text{MHz}}$  relation (orange line) and broken power-law parameterisation (dashed blue line) found by G18 (which was based on a sample identified based on the MPA-JHU catalogue at optical wavelengths). The apparent offset to larger  $L_{150\text{MHz}}$  in this sample relative to the G18 relations is an artefact of the data having been censored by the  $3\sigma$  threshold in  $L_{150\text{MHz}}$ , and it does not account for the large majority of sources in our sample that fall

<sup>6</sup> This is not unreasonable given the immense SED-fitting effort required to replicate them for the full IRAC-selected sample.



**Fig. 2.** Histograms (clockwise from the top-left) showing the SFR ( $\psi$ ), stellar mass, 150 MHz S/N, and redshift distributions for the 118 517 galaxies in our  $3.6\mu\text{m}$ -selected  $z < 1$  sample after applying the cuts discussed in Sect. 2. The overlaid orange histograms show the corresponding distributions for the subset of the sample that is detected to  $\geq 3\sigma$  at 150 MHz.

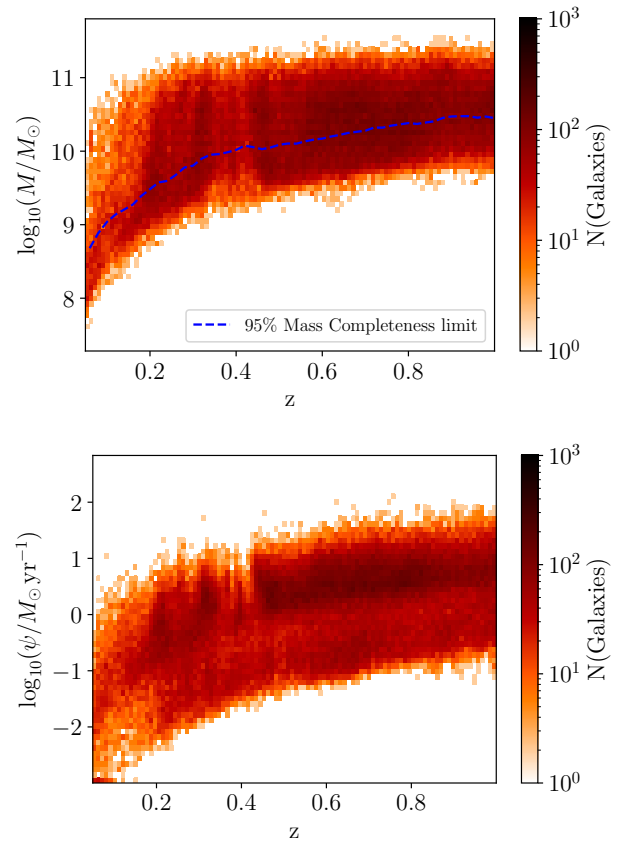
below it. The formally non-detected (i.e.  $< 3\sigma$ ) radio sources have huge potential diagnostic power due to their numerical dominance (e.g. Zwart et al. 2015; Malefahlo et al. 2020), but they are challenging to visualise in the SFR- $L_{150\text{MHz}}$  plane. As the lower-right panel of Fig. 2 shows, the vast majority of sources in our IRAC-selected sample have 150 MHz flux densities with  $< 3\sigma$  significance, and a substantial minority have negative measured flux densities. It is critical that we also account for these sources since failure to do so clearly truncates the luminosity distribution at a given SFR, and, left unchecked, this would have the potential to introduce significant bias into our results.

Finally, to test the possible influence of sample incompleteness on our results, we used the method from Pozzetti et al. (2010) to identify those galaxies with masses in excess of the 95% mass completeness limit as a function of redshift (shown as the dashed blue line in the upper panel of Fig. 3). We repeated all of the following analyses considering only those galaxies with masses in excess of the 95% completeness limit and found that our results did not change when the uncertainties were taken into account. We therefore conclude that incompleteness does not exert significant influence on our results.

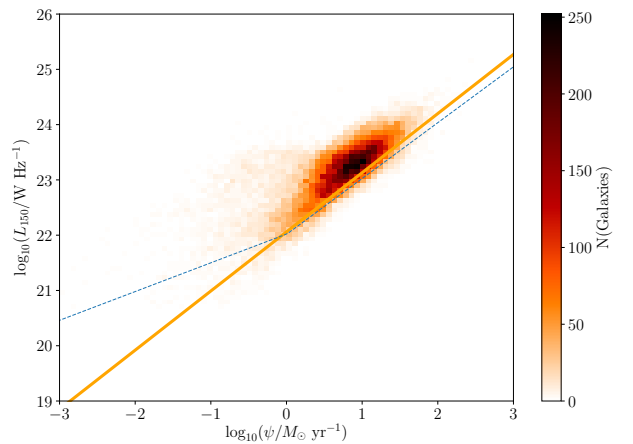
### 3. Results

#### 3.1. The SFR- $L_{150\text{MHz}}$ relation

We determined the SFR- $L_{150\text{MHz}}$  relation in ELAIS-N1, while accounting for uncertainties on both SFR and  $L_{150\text{MHz}}$ , as follows. First, for each source in our sample, we created a two-dimensional PDF in the SFR- $L_{150\text{MHz}}$  parameter space with logarithmic axes in both directions, with 70 equally spaced bins of SFR between  $-3 < \log_{10}(\psi/M_{\odot} \text{ yr}^{-1}) < 3$  and 180 equally log-spaced bins of  $L_{150\text{MHz}}$  between  $17 < \log_{10}(L_{150\text{MHz}}/\text{W Hz}^{-1}) < 26$ . We generated each source's PDF by creating a histogram using the aforementioned bins, populated by 100 samples each in

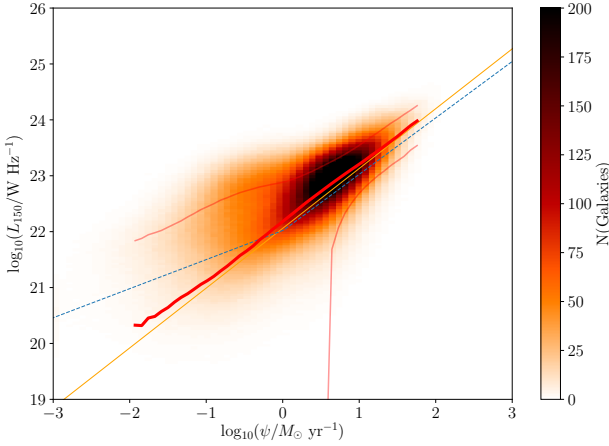


**Fig. 3.** Heatmaps showing the variation in median-likelihood stellar mass in  $M_{\odot}$  (top) and SFR ( $\psi$ ) in units of  $M_{\odot} \text{ yr}^{-1}$  (bottom) as a function of redshift in our IRAC-selected sample. The redshift axis is common to both plots, and the number of galaxies in each bin is indicated by the colour scales to the right. Upper panel: the dashed blue line indicates the 95% mass completeness limit as a function of redshift, derived using the method from Pozzetti et al. (2010).



**Fig. 4.** SFR- $L_{150\text{MHz}}$  plane populated by the 25 777 sources with  $\geq 3\sigma$  detections in the 150 MHz data, with the colour bar to the right indicating the bin occupancy. The best-fit relations from G18 are overlaid for comparison; the mass-independent relation is shown in orange, while the broken power-law relation evaluated at  $10^{10} M_{\odot}$  is shown as the dashed blue line.

the SFR and  $L_{150\text{MHz}}$  directions, assuming that the uncertainties in SFR and  $L_{150\text{MHz}}$  are uncorrelated. For the  $L_{150\text{MHz}}$  values, we adopted a normally distributed error distribution (in linear space) with median and standard deviation equal to those derived using



**Fig. 5.** Stacked heatmap showing the two-dimensional PDF for SFR and  $L_{150\text{MHz}}$  for all 118 517 galaxies in our sample, including the uncertainties on the SFR (derived based on the percentiles of the MAGPHYS SFR PDF) and on  $L_{150\text{MHz}}$  (using the pixel flux densities and rms values at the redshift in the Duncan et al. 2021 catalogue). The “effective” number of galaxies in each bin – we recall that each galaxy is sampled 100 times (with each sample representing 0.01 galaxies) and can therefore contribute to multiple pixels in the stack – is indicated by the colour bar to the right. The best-fit SFR- $L_{150\text{MHz}}$  relationship from G18 is overlaid as the solid orange line, along with the broken power-law G18 relation (which we have evaluated at a canonical stellar mass of  $10^{10} M_{\odot}$ , and which is shown as the dashed blue line) along with (in red) the 16th, 50th, and 84th percentiles of the  $L_{150\text{MHz}}$  distribution at each SFR.

the measured flux densities and an rms from the LoTSS data at each source’s redshift in our adopted cosmology. To ensure that the low signal to noise and negative values of  $L_{150\text{MHz}}$  are included in the PDF – essential to avoid biasing them by censure – we arbitrarily assigned samples with  $\log_{10} L_{150\text{MHz}} < 17$  (including the negative values) to the lowest bin of the  $L_{150\text{MHz}}$  PDF. For the SFRs, we adopted an asymmetric error distribution with a median equal to the median-likelihood value of the MAGPHYS SFR PDF, and a different standard deviation on either side of the median, equal to the difference between the 84th (16th) and 50th percentiles of the SFR PDF for the positive (negative) wings of the error distribution. We then summed these PDFs over the whole sample to arrive at a stacked PDF showing the distribution of SFR and  $L_{150\text{MHz}}$  that includes the uncertainties for the whole sample in both directions. This is shown as the heatmap in the background of Fig. 5, with the colour bar to the right indicating the number of galaxies in each bin<sup>7</sup>.

One of the most interesting results from G18 was the discovery of an upturn in the SFR- $L_{150\text{MHz}}$  relation at low SFRs, below  $\log_{10}(\psi/M_{\odot} \text{yr}^{-1}) \approx 0$ , and the G18 best-fit broken power-law relation is shown as the dashed blue line in Fig. 5 (for a stellar mass of  $10^{10} M_{\odot}$ , typical of galaxies in our sample). In order to identify the SFR- $L_{150\text{MHz}}$  relation in our data, we used a non-parametric approach. Non-parametric methods have the implicit advantage of being agnostic about the precise form of any relation that they may recover, and are therefore ideal for determining whether the data support an upturn at a low SFR, of the type seen in G18, or not. At each SFR, we calculated the 16th, 50th, and 84th percentiles of the  $L_{150\text{MHz}}$  distribution in the stacked two-dimensional PDF. The 50th percentile of the distribution then corresponds to our median-likelihood estimate of the SFR- $L_{150\text{MHz}}$  relation at that particular SFR, while the 16th and

84th percentiles also depend on the combination of the uncertainties on the luminosity estimates, plus any intrinsic scatter in the relation itself. We estimated the uncertainties on the median-likelihood value in each SFR bin by using the median statistics method from Gott et al. (2001). A further appealing feature of a median-likelihood estimate is that it has a degree of built-in resistance to outliers, such as might be expected from, for example, some residual minority of sources hosting unidentified radio excess due to AGN activity. We also note that deriving the best-fit relation in this way does not require us to account for the intrinsic dispersion of the relation itself,  $\sigma_L$ , to which we will return in Sect. 3.2.

The results are shown as the red lines overlaid on Fig. 5, with the thick line corresponding to the median-likelihood estimate of SFR- $L_{150\text{MHz}}$  over the range  $-2 < \log_{10}(\psi/M_{\odot} \text{yr}^{-1}) < 2$ , with the 16th and 84th percentiles shown as the thin red lines. It is immediately clear that the median-likelihood relation is similar to the mass-independent relation found by G18, albeit slightly flatter, and that the recovered values appear consistent with a power-law relationship across the full range of the SFRs spanned by our sample, with no evidence of an upturn in  $L_{150\text{MHz}}$  at low SFRs. To determine the best-fit parameters, we adopted the form of the SFR- $L_{150\text{MHz}}$  relation from G18, specifically

$$L_{150\text{MHz}} = L_1 \psi^{\beta}, \quad (1)$$

and find best-fit values of  $\log_{10} L_1 = 22.181 \pm 0.005$  and  $\beta = 1.041 \pm 0.007$ , where the uncertainties have been determined using the emcee (Foreman-Mackey et al. 2013) Monte Carlo Markov chain (MCMC) algorithm with 16 walkers and a chain length of 10 000 samples.

In Appendix C, we present a suite of simulations conducted to test how well we can recover a known SFR- $L_{150\text{MHz}}$  relation using this method, and we find that the best-fit estimates are likely to be systematically offset by  $\Delta \log_{10} L_1 \approx 0.040$  and  $\Delta \beta \approx 0.016$ . Correcting for these effects gives our best estimate of the overall SFR- $L_{150\text{MHz}}$  relation, which is  $\log_{10} L_{150\text{MHz}} = (22.221 \pm 0.008) + (1.058 \pm 0.007) \log_{10}(\psi/M_{\odot} \text{yr}^{-1})$ . Although these values are formally inconsistent with the results from G18 (i.e.  $\log_{10} L_1 = 22.06 \pm 0.01$ ,  $\beta = 1.07 \pm 0.01$ ), the fact that they are this close is encouraging, given the large differences in methodology and sample definition.

Figure 5 shows a clear offset between the apparent peak of the stacked PDF (background colour scale) and the median-likelihood values (thick red lines), which corresponds to our estimate of the SFR- $L_{150\text{MHz}}$  relation. This effect results from sampling linear  $L_{150\text{MHz}}$  values in a logarithmic PDF, which means that the lower half of the PDF is effectively spread out over a larger number of bins compared to the higher half<sup>8</sup>. This effect is also clearly apparent in the simulations of our method discussed in Appendix C, which underlines that the non-Gaussianity in the PDF does not stop the median-likelihood values (of the individual sources and of the population) from being able to recover the true SFR- $L_{150\text{MHz}}$  relation.

The third important thing to notice about Fig. 5 is that the 16th percentile of the  $L_{150\text{MHz}}$  distribution (shown as the lower red line) rapidly decreases below the bottom of the plotting window at  $\log_{10} \psi < 0.7$ . This highlights the importance of including the formally undetected sources in this study (i.e. those with 150 MHz flux density  $< 3\sigma$ ), and this importance of

<sup>7</sup> Appendix A describes an illustrative example of this method, along with more of the intermediate steps for the interested reader.

<sup>8</sup> We can consider, for example, that the  $\pm 2\sigma$  confidence interval in log values for a  $3\sigma$  source with a true luminosity of  $10^{24} \text{W Hz}^{-1}$  is  $23.52 < \log_{10}(L_{150\text{MHz}}/\text{W Hz}^{-1}) < 24.22$ . The range is more than twice as large on the negative side as on positive side.



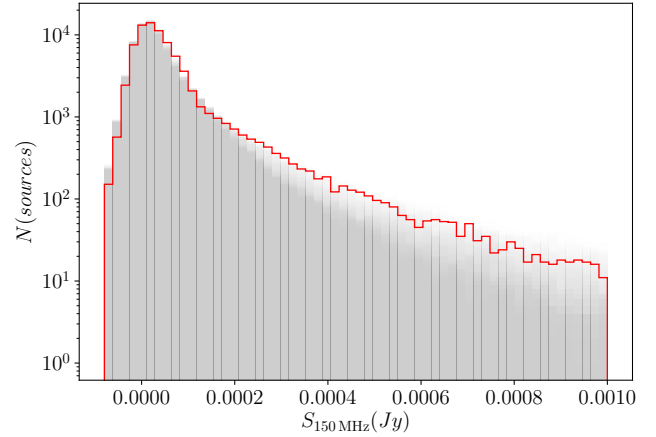
course increases as we move to lower SFRs, where an increasing fraction of objects have 150 MHz flux densities below  $1\sigma$ . The simulations discussed in Appendix C further highlight the importance of accounting for the negative  $L_{150\text{MHz}}$  samples since not doing so would introduce significant positive bias into the median flux density.

It is tempting to also determine the SFR- $L_{150\text{MHz}}$  relation by finding the “ridge-line” in the stacked two-dimensional PDF that corresponds to the modal value of  $L_{150\text{MHz}}$  at a given SFR. If we do this, we obtain good agreement with the above median-likelihood estimates at  $\log_{10}(\psi / M_{\odot} \text{yr}^{-1}) > 1$ ; however, as we move to lower SFRs, we begin to see positive bias introduced, similar to the upturn seen in G18. This is another facet of the aforementioned issue with increasingly sampling the noise distribution as we move to lower SFRs: if all sources were at the same redshift, then we are effectively attempting to plot a histogram of the logarithm of a normal distribution in luminosity that is centred very close to zero, and the ridge-line becomes increasingly biased.

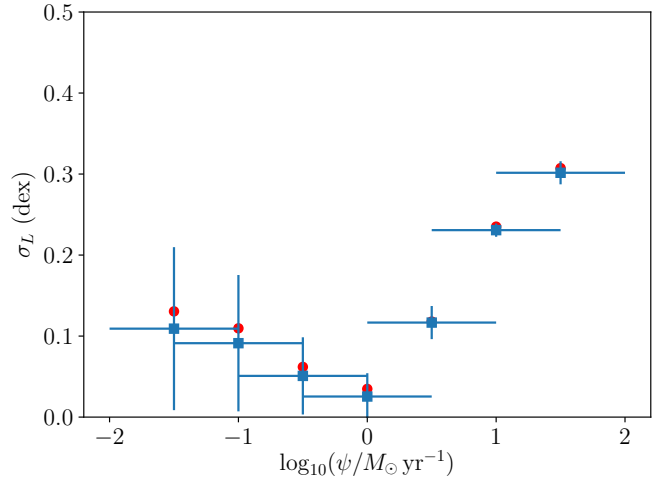
### 3.2. Scatter on SFR- $L_{150\text{MHz}}$

In addition to the form of the SFR- $L_{150\text{MHz}}$  relation, it is also of interest to determine the scatter on the relation itself,  $\sigma_L$ , which is usually quoted in logarithmic terms (dex). Whatever the cause,  $\sigma_L$  is important since it forms a key part of the process of identifying AGN in 150 MHz samples on the basis of an SFR-dependent radio excess (Best et al., in prep.). It is also of interest when studying the so-called “main sequence” of star formation – the relationship between the stellar mass and SFR of star-forming galaxies (e.g. Noeske et al. 2007; Johnston et al. 2015; Schreiber et al. 2015) – using radio observations (e.g. Leslie et al. 2020). The width of the main sequence has been interpreted as a manifestation of the variation in a star-forming galaxy’s gas supply (Tacchella et al. 2016); therefore, an additional source of scatter in the SFR indicator itself has the potential to bias the results if it is not measured and accounted for.

To measure the scatter on the SFR- $L_{150\text{MHz}}$  in our data, we used Monte Carlo simulations. We did this by creating multiple realisations of our sample, using the best-fit redshifts alongside random draws from the MAGPHYS SFRs (assuming the same asymmetric error distribution as in Sect. 3.1) and from the best-fit relation (Eq. (1)), with a scatter in the range  $0.0 < \sigma_L < 0.5$  dex, to calculate a model  $L_{150\text{MHz}}$  for each source. We used these values to derive “true” flux densities and simulate measurement errors based on realisations of the values in the rms map at the position of that source in the real data. For each simulation, we conducted two-sample Kolmogorov-Smirnov (KS) tests to compare the model flux densities with the real flux density distribution for any choice of SFR and redshift, to in turn determine the degree of support for the hypothesis that the simulated flux densities are consistent with being drawn from the same distribution as the real values, as a function of  $\sigma_L$ . Figure 6 shows the true pixel flux density distribution in red, overlaid on one set of Monte Carlo simulations covering the full range in scatter (shaded grey). The individual grey histograms are transparent, such that lighter grey regions reveal the variation in the flux density distribution for the range of scatter considered. Following Macfarlane et al. (2020), we truncated the distribution above a flux density of 1 mJy to avoid giving undue influence to residual undiagnosed AGN in our sample (this does not make a significant difference to the results). We determined the mean  $P$  returned by the KS test (averaging over all 10 000 Monte Carlo simulations) as a function of  $\sigma_L$  and marginalised the resulting



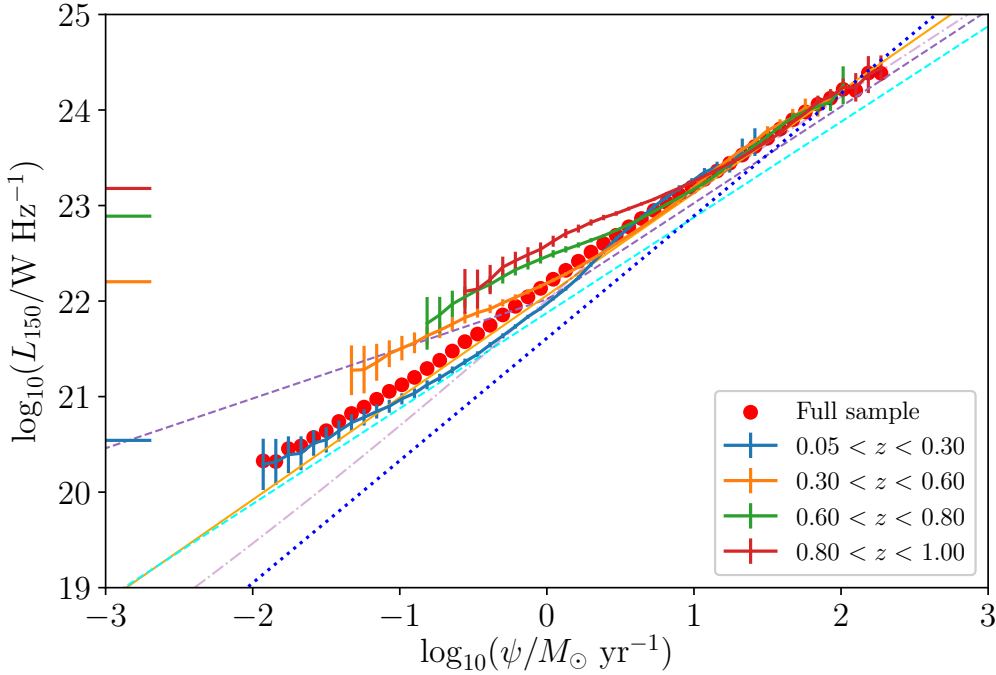
**Fig. 6.** Histogram with logarithmic ordinates, showing the true observed flux density distribution (in red) overlaid on individual Monte Carlo realisations (shown in grey with transparency, such that lighter shading indicates the range of outputs) based on assuming the best-fit SFR- $L_{150\text{MHz}}$  relation from Sect. 3.1 with the full range of intrinsic scatter  $0.0 < \sigma_L < 0.5$ .



**Fig. 7.** Variation in median-likelihood estimates of  $\sigma_L$  as a function of the central value of each SFR bin. The error bars in the vertical direction are derived from the 16th and 84th percentiles of the PDF, while in the horizontal direction the error bars indicate the bin width. Also overlaid are red circles, indicating the Bayesian estimates of scatter from the PDF, calculated according to the product of  $\sum \sigma_L P(\sigma_L)$ .

distribution. We were then able to calculate median-likelihood estimates of  $\sigma_L$  along with uncertainties by estimating the 16th, 50th, and 84th percentiles of the PDF, as well as Bayesian estimates calculated according to  $\sum \sigma_L P(\sigma_L)$ .

Figure 7 shows our results from using the KS tests to determine the level of support in the data for different values of  $\sigma_L$  as a function of SFR in bins with a constant width of 1 dex in SFR on a sliding scale from  $-2 < \log_{10}(\psi / M_{\odot} \text{yr}^{-1}) < 2$ . We only detect significant scatter about SFR- $L_{150\text{MHz}}$  at  $\log_{10}(\psi / M_{\odot} \text{yr}^{-1}) > 0$ , with  $\sigma_L$  apparently increasing with SFR and reaching  $0.31 \pm 0.01$  dex at  $1 < \log_{10}(\psi / M_{\odot} \text{yr}^{-1}) < 2$ . This may be because the scatter in the flux density distribution is increasingly dominated by the sensitivity of the LOFAR observations as we approach lower SFRs, rather than by the physical effect that is most visible at larger SFRs. It will be of great interest to see whether this effect persists with larger samples of  $\log_{10}(\psi / M_{\odot} \text{yr}^{-1}) < 0$  galaxies from the wider-area LoTSS



**Fig. 8.** Median-likelihood estimates of the SFR- $L_{150\text{MHz}}$  relation for the whole sample (red circles) and in four redshift bins ( $0.05 < z < 0.30$ ,  $0.30 < z < 0.60$ ,  $0.60 < z < 0.80$ , and  $0.80 < z < 1.00$  are shown as the blue, orange, green, and red crosses, respectively). Error bars on the median  $L_{150\text{MHz}}$  are calculated using the median statistics method from [Gott et al. \(2001\)](#). Also overlaid are the SFR- $L_{150\text{MHz}}$  relationships from [Bell \(2003\)](#), dot-dashed purple line) and [Murphy et al. \(2011\)](#), dashed cyan line), converted to 150 MHz assuming a canonical spectral index  $\alpha = 0.7$ , and the 150 MHz relations from G18 (mass-independent shown as the solid orange line, broken power law evaluated at  $10^{10} M_{\odot}$  as the dashed purple line) and [Wang et al. \(2019\)](#) (shown as the dotted blue line). The literature calibrations have been converted to our adopted IMF from [Chabrier \(2003\)](#) using the factors recommended in [Madau & Dickinson \(2014\)](#). The horizontal coloured lines immediately to the right of the left-hand vertical axis indicate the luminosity corresponding to the  $60 \mu\text{Jy}$  at the lower redshift bound of each bin, as indicated by the colour (see text for details).

second data release once it is available, and in due course with the huge increase in sensitivity that the SKA will provide.

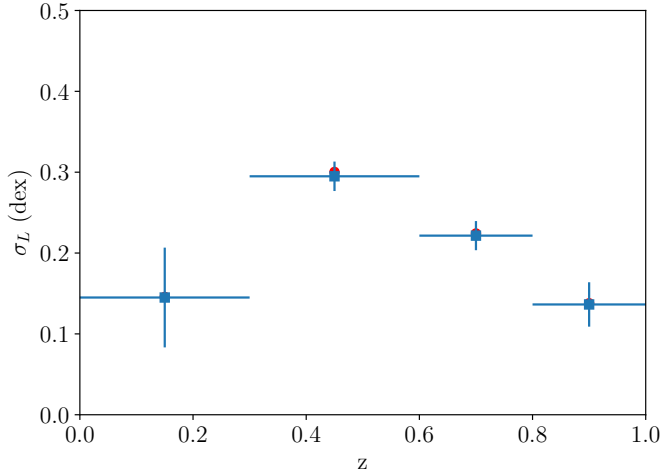
### 3.3. Evolution of SFR- $L_{150\text{MHz}}$ and $\sigma_L$

In order to investigate the possibility of evolution in the derived properties of SFR- $L_{150\text{MHz}}$ , we split the sample into four redshift bins and repeated the analysis from the previous two sections. Figure 8 shows the median-likelihood SFR- $L_{150\text{MHz}}$  derived over the full redshift range (in red) and overlaid with the values derived for each redshift bin (coloured as in the legend). For the purposes of comparison, we have once again overlaid the G18 relation (in orange) as well as the empirical relations from [Bell \(2003\)](#) and [Murphy et al. \(2011\)](#), both converted from the 1.4 GHz expectations assuming a canonical spectral index of  $\alpha = 0.7$  (similar to values in the literature, e.g. [Mauch et al. 2013](#); [Prescott et al. 2016](#)), and the relation from [Wang et al. \(2019\)](#). All of these literature relations have been converted to the [Chabrier \(2003\)](#) IMF used in this work where necessary, using the corrections provided in [Madau & Dickinson \(2014\)](#). To indicate the flux density scale in each redshift bin, we have also overlaid horizontal coloured lines adjacent to the left-hand vertical axis, which indicate values of  $L_{150\text{MHz}}$  corresponding to  $60 \mu\text{Jy}$  (i.e. the approximate  $3\sigma$  limit in the deepest regions of the 150 MHz data) at the lowest redshift bound of each bin (indicated by the colour of the line).

At  $\log_{10}(\psi/M_{\odot}\text{yr}^{-1}) > 1$ , there is negligible difference apparent between the SFR- $L_{150\text{MHz}}$  relation derived over the full redshift range and the values in each redshift bin. However, at  $\log_{10}(\psi/M_{\odot}\text{yr}^{-1}) < 1$ , there is the first piece of evidence of

variation and perhaps evidence of an excess in radio luminosity that increases towards the lowest SFR end of the data sampled in each redshift range, where the individual galaxies' 150 MHz flux densities are formally undetected. To determine whether this effect is real or instrumental, we conducted further simulations, repeating the analysis using a known input SFR- $L_{150\text{MHz}}$  relation, sampling the observed  $z$  and SFR distributions and uncertainties alongside a realistic model for mass-dependent AGN contamination following [Sabater et al. \(2019\)](#), and modelling the effects of noise in the 150 MHz flux densities by sampling from the real data set. The simulations are discussed in further detail in Appendix C. We were unable to reproduce variations in SFR- $L_{150\text{MHz}}$ , such as the possible upturn seen at  $\log_{10}(\psi/M_{\odot}\text{yr}^{-1}) < 1$  using a mass-independent SFR- $L_{150\text{MHz}}$  relation of the form given by Eq. (1). If this effect is real, it may point to possible mass dependence in SFR- $L_{150\text{MHz}}$  as found by G18. We will return to this topic in Sect. 3.4.

To make a better comparison with G18, who used a  $z < 0.3$  sample that was defined based on SDSS spectroscopy, we also derived best-fit parameters for the corresponding redshift range in our sample (our lowest redshift bin). We find best-fit parameters of  $\log_{10} L_1 = 22.14 \pm 0.01$  and  $\beta = 1.22 \pm 0.01$ , corrected for residual bias in the same way as in Sect. 3.1. Although our  $z < 0.3$   $L_1$  estimate is comparable, the value for  $\beta$  that we obtain in the lowest redshift bin is significantly steeper than that from G18, who found  $L_1 = 22.06 \pm 0.01$  and  $\beta = 1.07 \pm 0.01$ . As mentioned in Sect. 3.1, there are significant differences in the methodology used by the two works, especially the selection function, meaning that we are not necessarily comparing like with like.



**Fig. 9.** Variation in  $\sigma_L$  for galaxies with  $0.5 < \log_{10} \psi < 1.5$ ,  $10 < \log_{10}(M/M_{\odot}) < 11$  as a function of redshift – with error bars derived from the 16th and 84th percentiles of the PDF – and centred on the median-likelihood values. The error bars in the redshift (horizontal) direction indicate the bounds of each redshift bin.

We also searched for evidence of evolution in the scatter using the same method as in Sect. 3.2, but limiting the SFR range to  $0.5 < \log_{10}(\psi/M_{\odot} \text{ yr}^{-1}) < 1.5$  (to mitigate possible variation in scatter over SFR) and limiting the stellar mass range to  $10 < \log_{10}(M/M_{\odot}) < 11$  (to mitigate any possible influence of the mass dependence in SFR- $L_{150\text{MHz}}$  found by G18), using the same redshift bins as above. In Fig. 9, we plot the derived values as a function of redshift, along with their error bars. While some variation in  $\sigma_L$  is possible within the uncertainties, the data do not show any evidence for linear evolution in  $\sigma_L$ , at least out to  $z = 1$ .

### 3.4. Mass dependence of SFR- $L_{150\text{MHz}}$

Finally, we also considered the possibility of stellar mass dependence in SFR- $L_{150\text{MHz}}$  as discussed by G18 and Read (2019, and in prep.). To do this, we introduced a three-dimensional version of the method we used in Sect. 3.1, now producing 100 samples in each of the SFR, stellar mass, and  $L_{150\text{MHz}}$  directions to make a three-dimensional PDF for each source. Once again, we accounted for possible asymmetry in the SFR and stellar mass error bars by using the 16th, 50th, and 84th percentiles of the MAGPHYS estimates and sampling from a linear space in the uncertainty on  $L_{150\text{MHz}}$ . As before, we then summed over all 118 517 sources in our sample. We used 50 equally spaced logarithmic bins of stellar mass between  $7.5 < \log_{10}(M/M_{\odot}) < 11.8$ , 60 equally spaced logarithmic bins of SFR between  $-3 < \log_{10}(\psi/M_{\odot} \text{ yr}^{-1}) < 3$ , and 180 equally spaced logarithmic bins of  $L_{150\text{MHz}}$  between  $17 < \log_{10}(L_{150\text{MHz}}/\text{WHz}^{-1}) < 26$  to calculate the median-likelihood  $L_{150\text{MHz}}$  in each bin. The left-hand panel of Fig. 10 shows values of median-likelihood  $L_{150\text{MHz}}$  as a function of the SFR at a constant stellar mass indicated by the colour bar, effectively slicing through our three-dimensional stellar mass–SFR- $L_{150\text{MHz}}$  distribution. Similarly, the right-hand panel shows corresponding values of  $L_{150\text{MHz}}$  as a function of stellar mass at a fixed SFR (with the SFR indicated for each line by the colour bar). In both panels, we show only those bins populated by at least 15 galaxies (after accounting for the fact that each galaxy is sampled 100 times), though there are still some effects of small number statistics visible in the lowest SFR bins, particularly in the right-hand panel. Nevertheless, the left-hand

panel reveals a variation of at least 0.5 dex in  $L_{150\text{MHz}}$  for a given SFR, depending on the stellar mass, while the right-hand panel shows more than 2 dex of SFR-dependence in  $L_{150\text{MHz}}$  at a fixed stellar mass. These effects are clearly large enough to potentially account for the difference in SFR- $L_{150\text{MHz}}$  that we observe in our lowest redshift bin relative to G18, and which we discussed in Sect. 3.3.

To quantify our results, we used the mass-dependent parameterisation from G18:

$$L_{150} = L_C \psi^{\beta} \left( \frac{M}{10^{10} M_{\odot}} \right)^{\gamma} \quad (2)$$

and obtain best-fit values of  $\log_{10} L_C = 22.111 \pm 0.004$ ,  $\beta = 0.850 \pm 0.005$ , and  $\gamma = 0.402 \pm 0.005$ , where G18 obtained  $\log_{10} L_C = 22.13 \pm 0.01$ ,  $\beta = 0.77 \pm 0.01$ , and  $\gamma = 0.43 \pm 0.01$ . We conducted an additional set of simulations, which we discuss in Appendix C, to determine how well we are able to recover a known relation of the form given by Eq. (2). We find that, using this method, the best-fit estimates are likely to be offset by a residual bias of  $\Delta\beta = 0.053$ ,  $\Delta\log_{10} L_C = 0.107$ , and  $\Delta\gamma = -0.072$ , giving best estimates of  $\beta = 0.903 \pm 0.012$ ,  $\log_{10} L_C = 22.218 \pm 0.016$ , and  $\gamma = 0.332 \pm 0.037$  where we have propagated the uncertainties based on the systematic corrections by adding in quadrature with those values derived from our MCMC fitting.

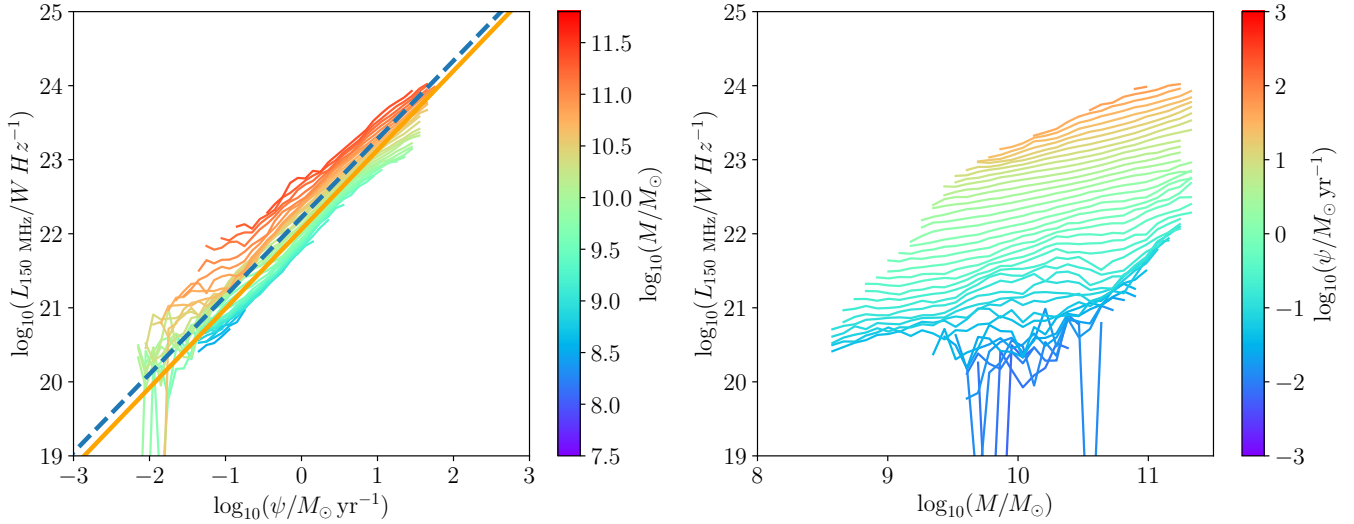
To visualise the improvement in accuracy resulting from Eq. (2), we show in Fig. 11 the same data as in Fig. 10, but with the mass and SFR dependency taken out (in the left- and right-hand panels, respectively) using the best-fit parameters from Eq. (2). Together, these plots reveal compelling evidence for mass dependence in SFR- $L_{150\text{MHz}}$ , in the sense that more massive galaxies have a larger radio luminosity at a fixed SFR. The fact that the mass dependence appears constant, and that it is clearly evident even in galaxies with stellar masses below  $10^{10} M_{\odot}$ , suggests that it is unlikely to be caused by some undiagnosed AGN contamination, given the mass dependence in AGN fraction shown by Sabater et al. (2019), for example.

Interestingly, we repeated the Monte Carlo simulations discussed in Sect. 3.2 to see whether assuming a mass-dependent SFR- $L_{150\text{MHz}}$  relation and including the MAGPHYS stellar mass information in the simulations made any difference to our measurements of  $\sigma_L$ . The results on both the SFR and redshift dependence in  $\sigma_L$  are consistent within the uncertainties whether we use a mass-dependent or mass-independent SFR- $L_{150\text{MHz}}$ , suggesting that mass dependence – at least as parameterised in Eq. (2) – cannot explain the scatter on SFR- $L_{150\text{MHz}}$ .

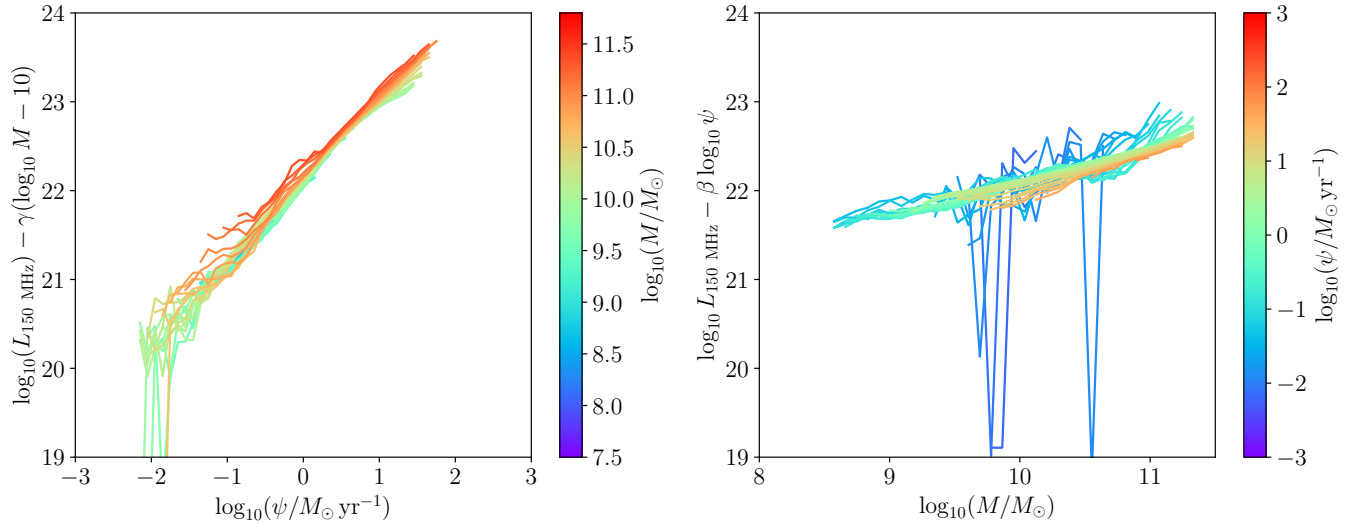
We also used a mass-dependent SFR- $L_{150\text{MHz}}$  relation as an input to the simulations discussed in Appendix C.1 and find that it allows us to recover an excess  $L_{150\text{MHz}}$  at low SFRs, very similar to the possible excess observed in the redshift bins of Fig. 8. Mass dependence on SFR- $L_{150\text{MHz}}$  may therefore be able to explain the possible variation revealed in Fig. 8, and it may also provide an explanation for the radio excess apparent at low SFRs first noticed by G18. We intend to revisit this issue in a future work.

## 4. Discussion and conclusions

We have studied the relationship between SFRs and 150 MHz luminosity using new sensitive deep field observations from the LoTSS Deep Fields DR1. Starting from a near-infrared selected sample, we leveraged the multi-wavelength aperture-matched



**Fig. 10.**  $L_{150\text{MHz}}$  dependence with SFR and stellar mass. *Left:* SFR- $L_{150\text{MHz}}$  relation as a function of stellar mass, indicated by the colour of the line and relative to the colour bar, overlaid on the mass-independent relation from G18 (solid orange line) and our best-fit estimate from Sect. 3.1 (dashed blue line). Each coloured line shows the median-likelihood  $L_{150\text{MHz}}$  at a given SFR and stellar mass. *Right:* relationship between stellar mass and median-likelihood  $L_{150\text{MHz}}$ , coloured as a function of SFR with the scale indicated by the bar to the right.



**Fig. 11.** Accounting for stellar mass effects in SFR- $L_{150\text{MHz}}$ . *Left:* SFR- $L_{150\text{MHz}}$  relationship with the mass dependence taken out using Eq. (2) and the best-fit parameters. Stellar mass is indicated by the colour bar to the right. *Right:* same as the left, but with the relationship between the stellar mass and radio luminosity normalised by the SFR using Eq. (2).

forced photometry and state-of-the-art photometric redshifts, alongside the new LoTSS maps of the ELAIS-N1 field, to produce stellar mass and SFR estimates using energy balance SED fits with the MAGPHYS package. We used 150 MHz flux densities from the ELAIS-N1 catalogue, plus pixel flux densities for the remaining 109 206 IRAC sources that are not identified as counterparts to catalogued 150 MHz sources, to estimate the median-likelihood  $L_{150\text{MHz}}$  as a function of the SFR and stellar mass.

The 150 MHz data used in this study are five times more sensitive than those used by G18, the sample size is eight times larger, and the multi-wavelength coverage in this field is far superior. This is true both in terms of depth (e.g. the HSC  $i$ -band data over ELAIS-N1 reach a  $5\sigma$  magnitude fainter than 26 mag, around four magnitudes deeper than the SDSS imaging used in G18) and in terms of the number of photometric bands that are available (Table 2 reveals that more than half of our sample have

$\geq 3\sigma$  flux densities in at least 16 bands, as compared to a maximum of 14 in G18). The LoTSS deep field data in ELAIS-N1 are around 35 times deeper than the FIRST data that have been used for previous studies of this topic (e.g. Hodge et al. 2008; Garn et al. 2009; Davies et al. 2017), assuming a standard canonical spectral index value of  $\alpha = 0.7$ ; they are comparable to the deepest degree-scale interferometric radio data in existence (e.g. Smolčić et al. 2017) but cover an area of sky five times larger.

Using a non-parametric approach, we find an apparently linear relationship between the SFR and  $L_{150\text{MHz}}$  over the range in SFR  $-2 < \log_{10}(\psi/M_{\odot}\text{yr}^{-1}) < 2$  of the form  $L_{150\text{MHz}} = L_1\psi^{\beta}$ , with best-fit parameters equal to  $\log_{10} L_1 = 22.221 \pm 0.008$  and  $\beta = 1.058 \pm 0.007$ . We find an SFR-dependent scatter about the SFR- $L_{150\text{MHz}}$  relation, reaching  $\sigma_L \approx 0.31 \pm 0.01$  dex at  $1 < \log_{10}(\psi/M_{\odot}\text{yr}^{-1}) < 2$ . Our inability to detect significant scatter at lower SFRs  $\log_{10}(\psi/M_{\odot}\text{yr}^{-1}) < 0$  may be a limitation of the sensitivity in the 150 MHz data, even though they are the

deepest in existence. Neither the scatter nor the high-SFR end ( $\log_{10}(\psi/M_{\odot}\text{yr}^{-1}) > 1$ ) of the best-fit relation show significant evidence for redshift evolution, with the latter found to be in agreement with, for example, Garn et al. (2009) and Duncan et al. (2020), out to larger SFRs and redshifts. Our results also agree with Calistro Rivera et al. (2017, who used a 150 MHz and *i*-band selected sample) out to our  $z < 1$  limit, though they do see evidence for evolution in  $\text{SFR-}L_{150\text{MHz}}$  in more distant sources.

The (close to) unitary slope that we determine for the best-fit  $\text{SFR-}L_{150\text{MHz}}$  relation is apparently consistent with expectations based on calorimetry models (e.g. Chi & Wolfendale 1990; Yun et al. 2001; Lacki et al. 2010), and it is similar to previous low-frequency results (e.g. Brown et al. 2017, who found  $\beta = 1.14 \pm 0.05$ , and G18, who found  $\beta = 1.07 \pm 0.01$ ). However, like G18 and Read (2019), we find a clear mass dependence in the  $\text{SFR-}L_{150\text{MHz}}$  relation, in the sense that higher mass galaxies have a larger  $L_{150\text{MHz}}$  at a fixed SFR. Our best-fit mass-dependent relation is  $\log_{10} L_{150\text{MHz}} = (0.90 \pm 0.01) \log_{10}(\psi/M_{\odot}\text{yr}^{-1}) + (0.33 \pm 0.04) \log_{10}(M/10^{10}M_{\odot}) + 22.22 \pm 0.02$ . Using a suite of realistic simulations, we have shown that the mass dependence can explain the possible observed deviation from linearity in the redshift-binned  $\text{SFR-}L_{150\text{MHz}}$  relation, as well as potentially the radio excess in low-SFR galaxies found by G18. This implies that the unitary slope we recovered in the overall  $\text{SFR-}L_{150\text{MHz}}$  may be a coincidence, and – assuming direct proportionality in the relationship between the far-infrared luminosity and the SFR as described, for example, in Kennicutt & Evans (2012) – we expect to observe similar mass dependence in the FIRC.

One possible explanation for these results could be a mass-dependent cosmic ray escape fraction that allows particles to remove energy from the galaxy before it can be radiated away at radio frequencies, especially in lower mass (smaller) galaxies. However, it is also important to consider whether undiagnosed AGN contamination can play a role since a radio excess that increases with growing stellar mass would be in keeping with previous results on the mass dependence of the AGN fraction (Sabater et al. 2019). However, the mass dependence apparent in our data is clear and consistent across the whole sample, including in galaxies with stellar masses below  $10^{10} M_{\odot}$ , implying that it is unlikely to be due to undiagnosed AGN contamination. At the same time, we are unable to rule out the presence of undiagnosed AGN in our sample altogether, especially low-excitation systems (which manifest as a large accretion-related radio excess but with little or no evidence for AGN in the multi-wavelength SED, e.g. Hardcastle et al. 2007; Best & Heckman 2012), but this type of undiagnosed AGN is unlikely to explain our results. Such issues relating to mass dependence are also discussed by Molnár et al. (2018) in the context of a variation in the FIRC in bulge-dominated galaxies (which also tend to be more massive than pure disk-dominated systems). We defer a more detailed discussion of this aspect of our results to a future work.

Irrespective of the cause, the size of this trend is such that failing to account for the stellar mass dependence can introduce systematic offsets on 150 MHz-derived SFRs, with a magnitude around 0.5 dex in either direction (consistent with the results shown in Read 2019), which are therefore potentially larger than the scatter inherent in  $\text{SFR-}L_{150\text{MHz}}$ . This value is also large enough that mass effects may explain the redshift evolution in the  $\text{SFR-}L_{150\text{MHz}}$  relation found by Calistro Rivera et al. (2017), although it is also possible that the evolution they report is partly due to their selection function (as underlined by our tests in Appendix B) and/or is only detectable at the higher redshifts probed by that work.

We continue to obtain new 150 MHz data over the ELAIS-N1 field, with the ultimate goal of reaching another factor of two greater sensitivity over the coming years. Complete optical spectroscopy using the *William Herschel* Telescope Enhanced Area Velocity Explorer (WEAVE: Dalton 2016) instrument will be obtained for every ELAIS-N1 150 MHz source brighter than  $100 \mu\text{Jy}$  as part of the WEAVE-LOFAR survey (Smith et al. 2016), which is scheduled to begin in the second quarter of 2021. Over five initial years of survey operations, WEAVE-LOFAR will obtain around a million spectra of LOFAR-selected sources in the best studied extragalactic fields in the Northern Hemisphere at every scale (ranging from deep fields such as ELAIS-N1, Lockman Hole, and Boötes and covering, for example, the whole of the H-ATLAS NGP field used by G18 and thousands of square degrees at high galactic latitudes), providing precise redshifts for virtually every source placed in a WEAVE fibre at  $z < 1$ . These new data will enable the use of extensive emission line classifications and Balmer-decrement derived SFRs to study the  $\text{SFR-}L_{150\text{MHz}}$  relation, which will in turn enable us to significantly improve on this work, including the full coverage of the luminosity-redshift plane simultaneously sampled by the wide and deep fields with highly uniform spectroscopy.

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## References

- Aihara, H., Arimoto, N., Armstrong, R., et al. 2018, *PASJ*, 70, S4  
 Antonucci, R. 1993, *ARA&A*, 31, 473  
 Appleton, P. N., Fadda, D. T., Marleau, F. R., et al. 2004, *ApJS*, 154, 147  
 Becker, R. H., White, R. L., & Helfand, D. J. 1995, *ApJ*, 450, 559

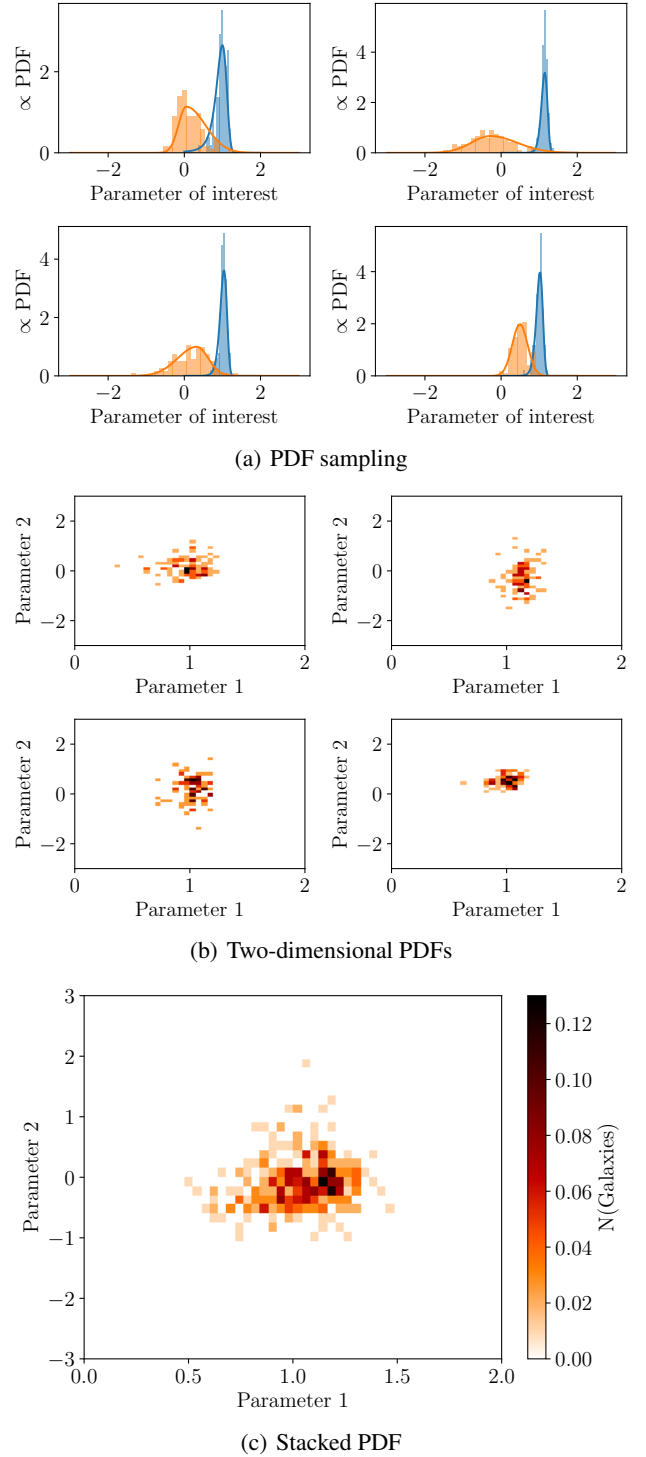
- Bell, E. F. 2003, *ApJ*, **586**, 794
- Berta, S., Lutz, D., Santini, P., et al. 2013, *A&A*, **551**, A100
- Best, P. N., & Heckman, T. M. 2012, *MNRAS*, **421**, 1569
- Bonato, M., Negrello, M., Mancuso, C., et al. 2017, *MNRAS*, **469**, 1912
- Bourne, N., Dunne, L., Ivison, R. J., et al. 2011, *MNRAS*, **410**, 1155
- Brinchmann, J., Charlot, S., White, S. D. M., et al. 2004, *MNRAS*, **351**, 1151
- Brown, M. J. I., Moustakas, J., Kennicutt, R. C., et al. 2017, *ApJ*, **847**, 136
- Bruzual, G., & Charlot, S. 2003, *MNRAS*, **344**, 1000
- Burgarella, D., Buat, V., & Iglesias-Páramo, J. 2005, *MNRAS*, **360**, 1413
- Calistro Rivera, G., Lusso, E., Hennawi, J. F., & Hogg, D. W. 2016, *ApJ*, **833**, 98
- Calistro Rivera, G., Williams, W. L., Hardcastle, M. J., et al. 2017, *MNRAS*, **469**, 3468
- Carilli, C. L., & Rawlings, S. 2004, *New Astron. Rev.*, **48**, 979
- Carnall, A. C., McLure, R. J., Dunlop, J. S., & Davé, R. 2018, *MNRAS*, **480**, 4379
- Chabrier, G. 2003, *PASP*, **115**, 763
- Chambers, K. C., Magnier, E. A., Metcalfe, N., et al. 2016, ArXiv e-prints [arXiv:1612.05560]
- Charlot, S., & Fall, S. M. 2000, *ApJ*, **539**, 718
- Chi, X., & Wolfendale, A. W. 1990, *MNRAS*, **245**, 101
- Condon, J. J. 1992, *ARA&A*, **30**, 575
- Cram, L., Hopkins, A., Mobasher, B., & Rowan-Robinson, M. 1998, *ApJ*, **507**, 155
- da Cunha, E., Charlot, S., & Elbaz, D. 2008, *MNRAS*, **388**, 1595
- Dalton, G. 2016, *ASP Conf. Ser.*, **507**, 97
- Davies, L. J. M., Huynh, M. T., Hopkins, A. M., et al. 2017, *MNRAS*, **466**, 2312
- De Gasperin, F., Williams, W. L., Best, P., et al. 2021, *A&A*, in press, <https://doi.org/10.1051/0004-6361/202140316>
- de Jong, T., Klein, U., Wielebinski, R., & Wunderlich, E. 1985, *A&A*, **147**, L6
- Delhaize, J., Smolčić, V., Delvecchio, I., et al. 2017, *A&A*, **602**, A4
- Dewdney, P. E., Hall, P. J., Schilizzi, R. T., & Lazio, T. J. L. W. 2009, *IEEE Proc.*, **97**, 1482
- Donley, J. L., Koekemoer, A. M., Brusa, M., et al. 2012, *ApJ*, **748**, 142
- Driver, S. P., Hill, D. T., Kelvin, L. S., et al. 2011, *MNRAS*, **413**, 971
- Dudzevičiūtė, U., Smail, I., Swinbank, A. M., et al. 2020, *MNRAS*, **494**, 3828
- Duncan, K. J., Sabater, J., Röttgering, H. J. A., et al. 2019, *A&A*, **622**, A3
- Duncan, K. J., Shivaee, I., Shapley, A. E., et al. 2020, *Phys. Rev. D*, **91**, 043510
- Duncan, K., Kondapally, R., Brown, M., et al. 2021, *A&A*, **648**, A4 (LoTSS SI)
- Flesch, E. W. 2019, ArXiv e-prints [arXiv:1912.05614]
- Foreman-Mackey, D., Hogg, D. W., Lang, D., & Goodman, J. 2013, *PASP*, **125**, 306
- Garn, T., Green, D. A., Riley, J. M., & Alexander, P. 2009, *MNRAS*, **397**, 1101
- Gott, J. R., I., Vogeley, M. S., Podariu, S., & Ratra, B. 2001, *ApJ*, **549**, 1
- Gürkan, G., Hardcastle, M. J., Smith, D. J. B., et al. 2018, *MNRAS*, **475**, 3010
- Haarsma, D. B., Partridge, R. B., Windhorst, R. A., & Richards, E. A. 2000, *ApJ*, **544**, 641
- Hardcastle, M. J., Evans, D. A., & Croston, J. H. 2007, *MNRAS*, **376**, 1849
- Hardcastle, M. J., Gürkan, G., van Weeren, R. J., et al. 2016, *MNRAS*, **462**, 1910
- Hardcastle, M. J., Croston, J. H., Shimwell, T. W., et al. 2019a, *MNRAS*, **488**, 3416
- Hardcastle, M. J., Williams, W. L., Best, P. N., et al. 2019b, *A&A*, **622**, A12
- Hayward, C. C., & Smith, D. J. B. 2015, *MNRAS*, **446**, 1512
- Helou, G., Soifer, B. T., & Rowan-Robinson, M. 1985, *ApJ*, **298**, L7
- Hodge, J. A., Becker, R. H., White, R. L., & de Vries, W. H. 2008, *AJ*, **136**, 1097
- Hopkins, A. M., Miller, C. J., Nichol, R. C., et al. 2003, *ApJ*, **599**, 971
- Hurley, P. D., Oliver, S., Betancourt, M., et al. 2017, *MNRAS*, **464**, 885
- Intema, H. T., Jagannathan, P., Mooley, K. P., & Frail, D. A. 2017, *A&A*, **598**, A78
- Ivison, R. J., Alexander, D. M., Biggs, A. D., et al. 2010a, *MNRAS*, **402**, 245
- Ivison, R. J., Magnelli, B., Ibar, E., et al. 2010b, *A&A*, **518**, L31
- Jarvis, M. J., Smith, D. J. B., Bonfield, D. G., et al. 2010, *MNRAS*, **409**, 92
- Johnston, R., Vaccari, M., Jarvis, M., et al. 2015, *MNRAS*, **453**, 2540
- Karim, A., Schinnerer, E., Martínez-Sansigre, A., et al. 2011, *ApJ*, **730**, 61
- Kellermann, K. I., & Owen, F. N. 1988, *Galactic and Extragalactic Radio Astronomy* (Berlin: Springer), 563
- Kennicutt, R. C., J. 1998, *ARA&A*, **36**, 189
- Kennicutt, R. C., & Evans, N. J. 2012, *ARA&A*, **50**, 531
- Kennicutt, R. C., J., Hao, C.-N., Calzetti, D., et al. 2009, *ApJ*, **703**, 1672
- Kondapally, R., Best, P., Hardcastle, M., et al. 2021, *A&A*, **648**, A3 (LoTSS SI)
- Lacki, B. C., & Thompson, T. A. 2010, *ApJ*, **717**, 196
- Lacki, B. C., Thompson, T. A., & Quataert, E. 2010, *ApJ*, **717**, 1
- Lacy, M., Storrer-Lombardi, L. J., Sajina, A., et al. 2004, *ApJS*, **154**, 166
- Lawrence, A., Warren, S. J., Almaini, O., et al. 2007, *MNRAS*, **379**, 1599
- Leslie, S., Schinnerer, E., Liu, D., et al. 2020, *ApJ*, **899**, 58
- Lisenfeld, U., Voelk, H. J., & Xu, C. 1996, *A&A*, **306**, 677
- Lofthouse, E. K., Kaviraj, S., Smith, D. J. B., & Hardcastle, M. J. 2018, *MNRAS*, **479**, 807
- Lonsdale, C. J., Smith, H. E., Rowan-Robinson, M., et al. 2003, *PASP*, **115**, 897
- Macfarlane, C., Best, P. N., Sabater, J., & Gürkan, G. 2020, *MNRAS*, submitted
- Madau, P., & Dickinson, M. 2014, *ARA&A*, **52**, 415
- Magnelli, B., Ivison, R. J., Lutz, D., et al. 2015, *A&A*, **573**, A45
- Mahatma, V. H., Hardcastle, M. J., Williams, W. L., et al. 2019, *A&A*, **622**, A13
- Malefahlo, E., Santos, M. G., Jarvis, M. J., White, S. V., & Zwart, J. T. L. 2020, *MNRAS*, **492**, 5297
- Martínez-Sansigre, A., Rawlings, S., Lacy, M., et al. 2005, *Nature*, **436**, 666
- Mauch, T., Klöckner, H.-R., Rawlings, S., et al. 2013, *MNRAS*, **435**, 650
- Mauduit, J. C., Lacy, M., Farrah, D., et al. 2012, *PASP*, **124**, 714
- McAlpine, K., Smith, D. J. B., Jarvis, M. J., Bonfield, D. G., & Fleuren, S. 2012, *MNRAS*, **423**, 132
- Molnár, D. C., Sargent, M. T., Delhaize, J., et al. 2018, *MNRAS*, **475**, 827
- Murphy, E. J. 2009, *ApJ*, **706**, 482
- Murphy, E. J., Helou, G., Kenney, J. D. P., Armus, L., & Braun, R. 2008, *ApJ*, **678**, 828
- Murphy, E. J., Condon, J. J., Schinnerer, E., et al. 2011, *ApJ*, **737**, 67
- Muzzin, A., Wilson, G., Yee, H. K. C., et al. 2009, *ApJ*, **698**, 1934
- Nisbet, D. M. 2018, PhD thesis, University of Edinburgh, UK
- Noeske, K. G., Weiner, B. J., Faber, S. M., et al. 2007, *ApJ*, **660**, L43
- Novak, M., Smolčić, V., Delhaize, J., et al. 2017, *A&A*, **602**, A5
- Oliver, S., Bock, J., Altieri, B., et al. 2012, *MNRAS*, **424**, 1614
- Pannella, M., Carilli, C. L., Daddi, E., et al. 2009, *ApJ*, **698**, L116
- Pannella, M., Elbaz, D., Daddi, E., et al. 2015, *ApJ*, **807**, 141
- Pearson, W. J., Wang, L., van der Tak, F. F. S., et al. 2017, *A&A*, **603**, A102
- Pozzetti, L., Bolzonella, M., Zucca, E., et al. 2010, *A&A*, **523**, A13
- Prescott, M., Mauch, T., Jarvis, M. J., et al. 2016, *MNRAS*, **457**, 730
- Read, S. C. 2019, PhD thesis, University of Hertfordshire, UK
- Read, S. C., Smith, D. J. B., Gürkan, G., et al. 2018, *MNRAS*, **480**, 5625
- Read, S. C., Smith, D. J. B., Jarvis, M. J., & Gürkan, G. 2020, *MNRAS*, **492**, 3940
- Retana-Montenegro, E., Röttgering, H. J. A., Shimwell, T. W., et al. 2018, *A&A*, **620**, A74
- Röttgering, H., Afonso, J., Barthel, P., et al. 2011, *JApA*, **32**, 557
- Sabater, J., Best, P. N., Hardcastle, M. J., et al. 2019, *A&A*, **622**, A17
- Sabater, J., Best, P., Tasse, C., et al. 2021, *A&A*, **648**, A2 (LoTSS SI)
- Sadler, E. M., Jackson, C. A., Cannon, R. D., et al. 2002, *MNRAS*, **329**, 227
- Sargent, M. T., Schinnerer, E., Murphy, E., et al. 2010, *ApJ*, **714**, L190
- Schreiber, C., Pannella, M., Elbaz, D., et al. 2015, *A&A*, **575**, A74
- Shimwell, T. W., Röttgering, H. J. A., Best, P. N., et al. 2017, *A&A*, **598**, A104
- Shimwell, T. W., Tasse, C., Hardcastle, M. J., et al. 2019, *A&A*, **622**, A1
- Shirley, R., Roehly, Y., Hurley, P. D., et al. 2019, *MNRAS*, **490**, 634
- Smith, D. J. B., & Hayward, C. C. 2018, *MNRAS*, **476**, 1705
- Smith, D. J. B., Dunne, L., Maddox, S. J., et al. 2011, *MNRAS*, **416**, 857
- Smith, D. J. B., Dunne, L., da Cunha, E., et al. 2012, *MNRAS*, **427**, 703
- Smith, D. J. B., Jarvis, M. J., Hardcastle, M. J., et al. 2014, *MNRAS*, **445**, 2232
- Smith, D. J. B., Best, P. N., Duncan, K. J., et al. 2016, *SF2A-2016: Proceedings of the Annual meeting of the French Society of Astronomy and Astrophysics*, eds. C. Reylé, J. Richard, L. Cambrésy, et al., 271
- Smolčić, V., Novak, M., Bondi, M., et al. 2017, *A&A*, **602**, A1
- Sudoh, T., Linden, T., & Beacom, J. F. 2020, ArXiv e-prints [arXiv:2005.08982]
- Sutherland, W., & Saunders, W. 1992, *MNRAS*, **259**, 413
- Tabatabaei, F. S., Schinnerer, E., Krause, M., et al. 2017, *ApJ*, **836**, 185
- Tacchella, S., Dekel, A., Carollo, C. M., et al. 2016, *MNRAS*, **457**, 2790
- Tasse, C., Shimwell, T., Hardcastle, M., et al. 2021, *A&A*, **648**, A1 (LoTSS SI)
- Upjohn, J. E., Brown, M. J. I., Hopkins, A. M., & Bonne, N. J. 2019, *PASA*, **36**, e012
- van der Kruit, P. C. 1971, *A&A*, **15**, 110
- van Haarlem, M. P., Wise, M. W., Gunst, A. W., et al. 2013, *A&A*, **556**, A2
- Wang, L., Rowan-Robinson, M., Norberg, P., Heinis, S., & Han, J. 2014, *MNRAS*, **442**, 2739
- Wang, L., Gao, F., Duncan, K. J., et al. 2019, *A&A*, **631**, A109
- Williams, W. L., Hardcastle, M. J., Best, P. N., et al. 2019, *A&A*, **622**, A2
- Wilson, G., Muzzin, A., Yee, H. K. C., et al. 2009, *ApJ*, **698**, 1943
- York, D. G., Adelman, J., Anderson, J. E., J., et al. 2000, *AJ*, **120**, 1579
- Yun, M. S., Reddy, N. A., & Condon, J. J. 2001, *ApJ*, **554**, 803
- Zwart, J. T. L., Jarvis, M. J., Deane, R. P., et al. 2014, *MNRAS*, **439**, 1459
- Zwart, J. T. L., Santos, M., & Jarvis, M. J. 2015, *MNRAS*, **453**, 1740

## Appendix A: Generating stacked PDFs by sampling

To illustrate the way that we created stacked two-dimensional PDFs using random sampling, we have included the following simple example based on two model parameters of interest. We assumed that both parameters have asymmetric uncertainties, but for different reasons. In the case of our first parameter (hereafter “parameter 1”), we assumed uncertainties that are normally distributed in linear space, but we wished to stack the PDF in log space (as in the case for  $L_{150\text{MHz}}$  in our real data set). For our second parameter (“parameter 2”), we proceeded as we did for our MAGPHYS estimates of stellar masses or SFRs since the MAGPHYS PDFs generally do not have an analytic form. We therefore made the simplifying assumption that the underlying PDF is Gaussian distributed, and we set the width of the distribution on the positive (negative) side using the difference between the 84th (16th) and 50th percentiles of the MAGPHYS PDF.

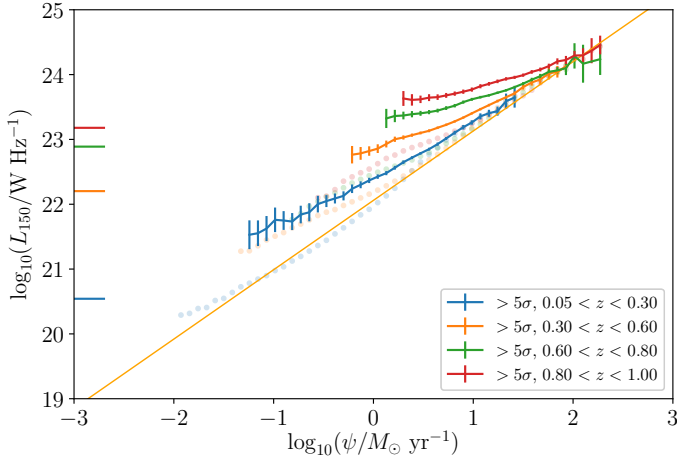
The four panels of Fig. A.1a show the PDFs assumed for parameters 1 and 2 in blue and orange, respectively, for some arbitrarily chosen values as well as for the hypothetical case of the four galaxies that we wish to stack. We then generated 100 samples from the assumed PDF for each parameter and for each galaxy; histograms of these samples have been overlaid in Fig. A.1a. Since we assume that the uncertainties are independent, we can use these samples to create a two-dimensional PDF for each object by creating a two-dimensional histogram using the same samples and then normalising. Examples of these two-dimensional PDFs for the four hypothetical objects are displayed in Fig. A.1b.

To study the relation between the two parameters for the full population (in this case, the hypothetical population of four galaxies, though in Sect. 3.1 we used around 120 000 galaxies), we can then stack the PDFs by summing up the values of the individual two-dimensional PDFs in each bin. The results for our hypothetical data set are shown in Fig. A.1c: with increasingly large numbers of galaxies, these PDFs become increasingly smooth, to the point that (as in Fig. 5) the distribution appears continuous despite the individual galaxies being sampled only 100 times. Each pixel in the stack shows how many galaxies we would expect to find in that bin; since we have sampled each galaxy multiple times and re-normalised (to retain the correct total number of galaxies), these need not be integers, as shown in Fig. A.1c.



**Fig. A.1.** Constructing PDFs for individual galaxies. *Top four panels:* analytic PDFs assumed for parameter 1 (blue) and parameter 2 (orange). We then created 100 random samples drawn from each distribution, and the results for each parameter are shown as the shaded histograms of the corresponding colours. *Middle panels:* two-dimensional PDFs for each hypothetical “galaxy”, derived by assuming that the samples shown in the top panels for each parameter are independent. *Bottom panel:* stack of the individual two-dimensional PDFs for the four model galaxies and the two indicative parameters shown in the upper panels. The colour bar indicates the number of model galaxies that we expect in each bin; the total obtained by summing all of the pixel values equals the total number of galaxies in the stack. Clearly, individual pixels can take non-integer values since we have sampled each galaxy 100 times.

## Appendix B: Studying SFR- $L_{150}$ with a 150 MHz-selected sample



**Fig. B.1.** SFR- $L_{150\text{MHz}}$  plane, as in Fig. 5. The coloured lines with error bars indicate the SFR- $L_{150\text{MHz}}$  relation recovered in each of the four redshift bins detailed in the legend – identical to those used in Fig. 8, and which are also reproduced here as the circles – but including only those sources with  $\geq 5\sigma$  150 MHz detections. Also overlaid is the best-fit relation from G18; the horizontal bars adjacent to the left-hand vertical axis indicate the luminosity corresponding to  $60\mu\text{Jy}$  at the lower bound of each redshift bin.

Many previous works have investigated the SFR – radio luminosity relation using a sample identified at radio frequencies (e.g. Bell 2003; Murphy et al. 2011; Brown et al. 2017; Calistro Rivera et al. 2017; Wang et al. 2019). To see the possible impact of this on our results for SFR- $L_{150\text{MHz}}$ , we repeated our analysis, but instead included only those sources that are detected at 150 MHz with  $\geq 5\sigma$  significance.

Figure B.1 reveals that making this kind of selection gives results that are biased to higher  $L_{150}$  at a given SFR (relative to both our IRAC-selected sample shown in Fig. 8 and to the best-fit relation from G18, which is shown as the solid orange line). Dividing such a sample into four redshift bins also reveals an apparent evolution in SFR- $L_{150\text{MHz}}$  – in the sense of an apparent increase in  $L_{150\text{MHz}}$  at a given SFR at higher redshift – which is not recovered when the full IRAC-selected sample is considered.

We conclude that great care is required when interpreting the results of this type of study, based on radio-frequency selected samples, even when using the most sensitive radio data in existence, such as the data used in this work from the LoTSS ELAIS-N1 deep field.

## Appendix C: Supporting simulations

In order to test that our method of determining the median-likelihood SFR is able to recover the true SFR- $L_{150\text{MHz}}$  relation, we conducted a set of simulations, each based on sampling 120 000 model galaxies with a range of SFRs, stellar masses, and redshifts from our real data. We assigned each galaxy a true  $L_{150\text{MHz}}$  value using an arbitrary mass-dependent SFR- $L_{150\text{MHz}}$  relation following Eq. (2), and we simulated Gaussian scatter about that relation assuming a standard width (in dex). Using Eq. (2), we were also able to simulate a stellar-mass independent SFR- $L_{150\text{MHz}}$  relation by setting  $\gamma = 0$ .

We then generated a set of mock observed data by converting the noiseless values of  $L_{150\text{MHz}}$  to true flux densities at 150 MHz,

**Table C.1.** Probability that a galaxy of a given stellar mass is assigned a radio luminosity excess in our simulations, based on the luminosity-averaged results from Sabater et al. (2019).

Mass range	$P(\text{AGN})$
$\log_{10}(M/M_{\odot}) < 10.00$	0.00
$10.00 < \log_{10}(M/M_{\odot}) < 10.50$	0.01
$10.50 < \log_{10}(M/M_{\odot}) < 10.75$	0.02
$10.75 < \log_{10}(M/M_{\odot}) < 11.00$	0.03
$11.00 < \log_{10}(M/M_{\odot}) < 11.25$	0.04
$11.25 < \log_{10}(M/M_{\odot}) < 11.50$	0.10
$11.50 < \log_{10}(M/M_{\odot}) < 12.00$	0.30

before adding on Gaussian noise using a random number generator multiplied by the flux density uncertainty measured for the real sources in our LoTSS catalogue. We also simulated measuring the SFRs (and stellar masses if required) by resampling the true values to give model-observed values with uncertainties that were sampled in the same way as our real MAGPHYS results. Our simulations account for AGN contamination by randomly assigning galaxies an excess radio luminosity drawn from a log-normal distribution with a mean of 23.86 and a standard deviation of 0.91, based on the mean and standard deviation of the  $L_{150\text{MHz}}$  of the flagged AGN in the Best et al. (in prep.) sample. The probability that a model galaxy is given such an excess luminosity is based on a mass-dependent probability using the luminosity-averaged results from Sabater et al. (2019), which are shown in Table C.1.

We then attempted to recover the known relation using the method discussed in Sect. 3.1. Figure C.1 shows an example visualisation of the two-dimensional PDF obtained from one of the simulations, assuming an SFR- $L_{150\text{MHz}}$  relation of the mass-independent form given in Eq. (1) with  $\beta = 1.01$ ,  $\log_{10} L_1 = 22.15$  (shown as the dashed pink line), and with scatter  $\sigma_L = 0.25$  dex. The thick red line shows the median-likelihood estimate of the  $L_{150\text{MHz}}$  in bins of SFR, while the purple line shows the best fit to the data over the range  $-1 < \psi < 1$ , the same range used for the real data. It is clear that, just as in the real data, the median-likelihood estimate (thick red line) appears to be offset to lower SFRs than the apparent peak in the PDF, and it suggests that our explanation for this effect in Sect. 3.1 is plausible.

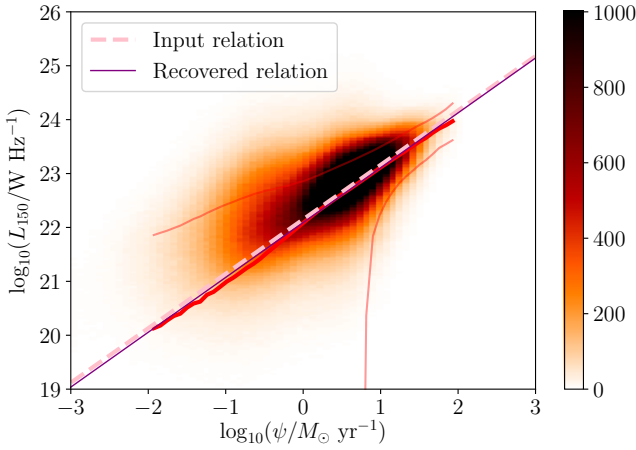
To better quantify the level of agreement between the input and output parameters in plausible realistic circumstances, we conducted 1000 Monte Carlo simulations of 120 000 model sources. We assumed a range of true values,  $0.5 \leq \beta \leq 1.0$ ,  $21.5 \leq \log_{10} L_C \leq 22.5$ , and  $0.30 < \gamma < 0.60$ , in ten equal steps each, and a fixed scatter of  $\sigma_L = 0.25$ .

We conducted the simulations twice: once simulating the recovery of the true input parameters for the mass-independent SFR- $L_{150\text{MHz}}$  (results displayed in Fig. C.2) and once simulating the recovery of the mass-dependent version (results displayed in Fig. C.3). In Fig. C.2, the left-hand panel shows the recovery of  $\beta$ , while the right-hand panel shows how well  $L_1$  is recovered. The 1:1 relation is indicated by the dotted line, while the best-fit relation between the input and the “observed” values is shown as the dashed grey line. The best fit relationships between the input and observed values are:

$$\beta^{\text{true}} = 1.152 \beta^{\text{obs}} - 0.142, \quad (\text{C.1})$$

$$\log_{10} \left( \frac{L_1^{\text{true}}}{10^{22} \text{WHz}^{-1}} \right) = 0.936 \log_{10} \left( \frac{L_1^{\text{obs}}}{10^{22} \text{WHz}^{-1}} \right) + 0.051, \quad (\text{C.2})$$



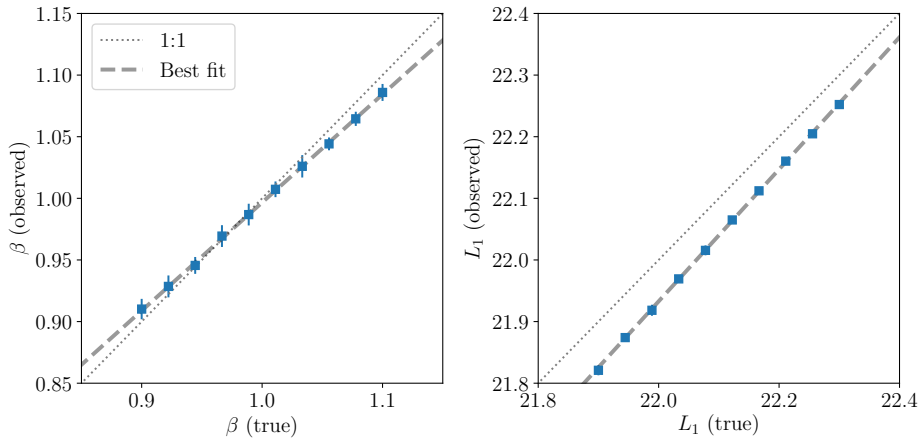


**Fig. C.1.** Heatmap showing the stacked two-dimensional PDF derived based on 120 000 galaxies with a redshift distribution sampled from our real data set and accounting for the uncertainties in both SFR and  $L_{150\text{MHz}}$  as in Fig. 5. The median-likelihood estimate of the SFR- $L_{150\text{MHz}}$  relation is shown as the thick red line, while the recovered relation (in purple) is very close to the true SFR- $L_{150\text{MHz}}$  relation (dashed pink line). The colour bar to the right shows the effective number of galaxies in each bin.

where the superscripts indicate the observed and true (i.e. model input) parameters. For the best-fit values quoted in Sect. 3.1, the difference in the parameters is small:  $\Delta\beta \equiv \beta^{\text{true}} - \beta^{\text{obs}} = 0.016$  and  $\Delta \log_{10} L_1 \equiv \log_{10} L_1^{\text{true}} - \log_{10} L_1^{\text{obs}} = 0.040$ , with scatter about the best-fit relation of  $\sigma_\beta = 0.006$  and  $\sigma_{\log_{10} L_1} = 0.006$ ; we quoted these systematic offsets and propagated the uncertainties on the values quoted in Sect. 3.1.

Similarly for the three-dimensional simulations, using ten bins of  $\beta$ ,  $\log_{10} L_C$ , and  $\gamma$ , the best-fit relations are:

$$\beta^{\text{true}} = 1.258 \beta^{\text{obs}} - 0.167, \quad (\text{C.3})$$



**Fig. C.2.** Simulations showing the observed values of  $\beta$  and  $L_1$  obtained using the method described in Sect. 3.1 as a function of the known input values, assuming a fixed width of  $\sigma_L = 0.25$  and mass-dependent AGN contamination following Sabater et al. (2019). The best-fit relations between the two sets of values – shown as the dashed grey line – are given in the text. The 1:1 line is shown as the dotted line.

$$\log_{10} \left( \frac{L_C^{\text{true}}}{10^{22} \text{WHz}^{-1}} \right) = 0.931 \log_{10} \left( \frac{L_C^{\text{obs}}}{10^{22} \text{WHz}^{-1}} \right) + 0.112, \quad (\text{C.4})$$

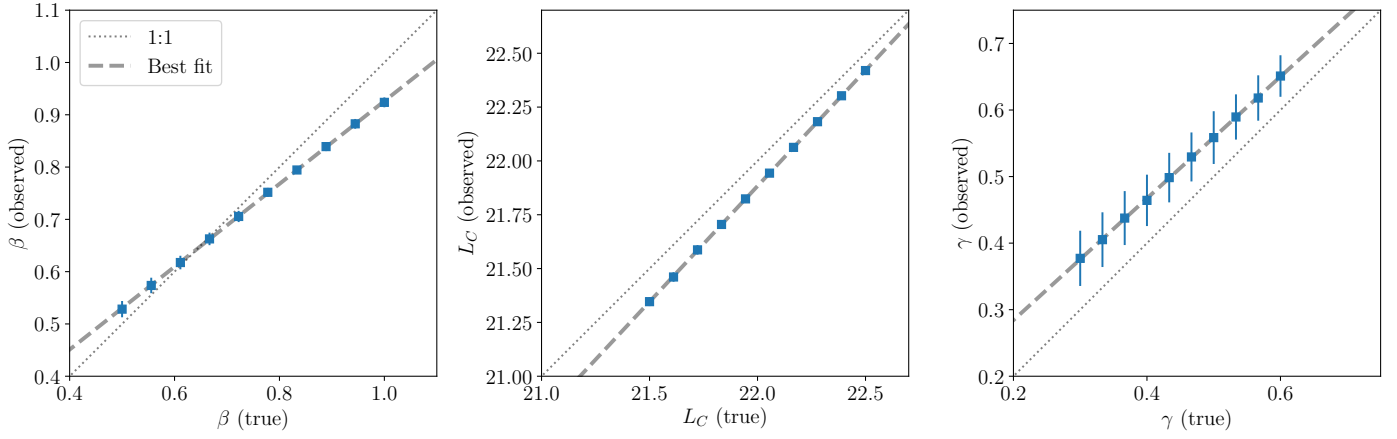
$$\gamma^{\text{true}} = 1.090 \gamma^{\text{obs}} - 0.108. \quad (\text{C.5})$$

For the best-fit values quoted in Sect. 3.4, the difference in the quoted parameters is  $\Delta\beta \equiv \beta^{\text{true}} - \beta^{\text{obs}} = 0.051$ ,  $\Delta \log_{10} L_C \equiv \log_{10} L_C^{\text{true}} - \log_{10} L_C^{\text{obs}} = 0.104$ , and  $\Delta\gamma \equiv \gamma^{\text{true}} - \gamma^{\text{obs}} = -0.057$ . The mean scatter about each best-fit relation dominates the uncertainties on these corrections, and the values are  $\sigma_\beta = 0.011$ ,  $\sigma_{\log_{10} L_C} = 0.016$ , and  $\sigma_\gamma = 0.037$ .

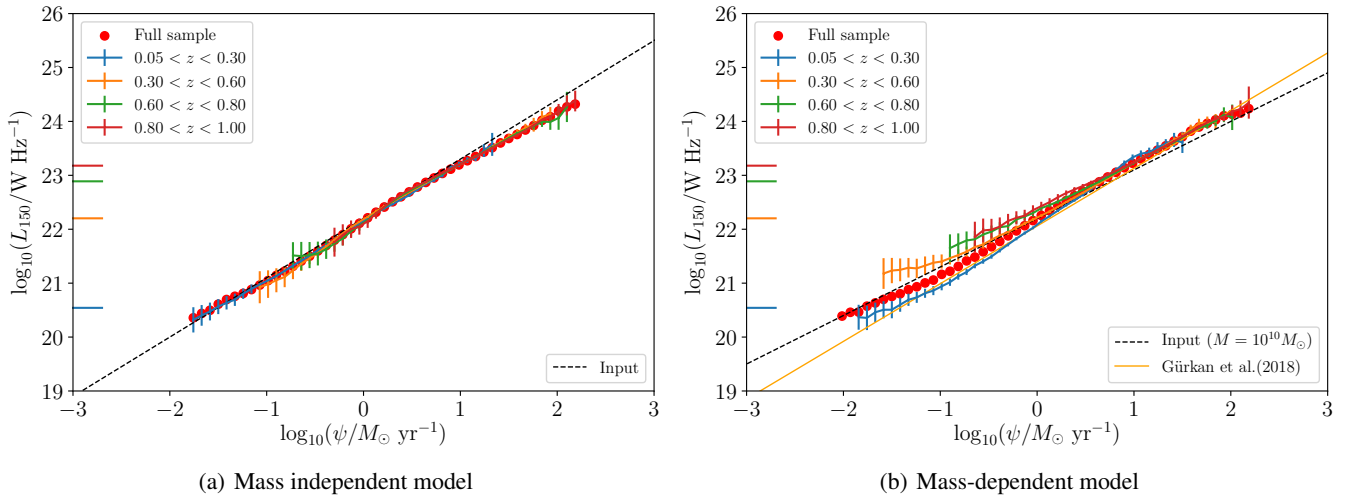
### C.1. Flux density limit or mass dependence?

Finally, we performed two further tests to see whether the possible redshift-dependent upturn to higher  $L_{150\text{MHz}}$  at lower SFRs shown in Fig. 8 is real or an artefact of the finite S/N available in the 150 MHz data set. First, we created a simulation identical to the ones in the previous section, with an input SFR- $L_{150\text{MHz}}$  relation that is independent of mass and repeated the analysis from Sect. 3.3 to see how well it was recovered. The left-hand panel of Fig. C.4 shows the results: while there is a slight bias towards a flatter SFR- $L_{150\text{MHz}}$  relation (as discussed in Appendix C), there is no evidence of an upturn, including below the approximate flux density limits indicated for each redshift bin by the horizontal error bars adjacent to the left vertical axis.

In the right-hand panel of Fig. C.4, we show the results of repeating an identical analysis using a second simulation, which includes a mass-dependent SFR- $L_{150\text{MHz}}$  relation of the form given in Eq. (2). The upturn to higher  $L_{150\text{MHz}}$  towards lower SFRs is now clear, underlining our view that this effect – if real – is consistent with being a consequence of the mass dependence in the SFR- $L_{150\text{MHz}}$  relation.



**Fig. C.3.** Simulations showing the observed values of  $\beta$  and  $L_1$  obtained using the method described in Sect. 3.1 as a function of the known input values, assuming a fixed width of  $\sigma_L = 0.25$  and AGN contamination following Sabater et al. (2019). The best-fit relations between the two sets of values are shown as the dashed grey lines, and they are parameterised as in the text. The dotted line in each panel shows the 1:1 relation.



**Fig. C.4.** Simulations of possible redshift evolution in the SFR- $L_{150\text{MHz}}$  relation using: (a) a mass-independent input relation of the form given in Eq. (1), and (b) a mass-dependent input relation of the form given in Eq. (2). The method used to obtain these figures is identical, and only the mass-dependent simulation reveals an upturn in  $L_{150\text{MHz}}$  at lower SFRs, as seen for the real data set analysed in Fig. 8 (see that caption for further details).