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# Probing Hot Gas Components of the Circumgalactic Medium in Cosmological Simulations with the Thermal Sunyaev–Zel’dovich Effect

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## Abstract

The thermal Sunyaev–Zel’dovich (tSZ) effect is a powerful tool with the potential for constraining directly the properties of the hot gas that dominates dark matter halos because it measures pressure and thus thermal energy density. Studying this hot component of the circumgalactic medium (CGM) is important because it is strongly impacted by star formation and active galactic nucleus (AGN) activity in galaxies, participating in the feedback loop that regulates star and black hole mass growth in galaxies. We study the tSZ effect across a wide halo-mass range using three cosmological hydrodynamical simulations: Illustris-TNG, EAGLE, and FIRE-2. Specifically, we present the scaling relation between the tSZ signal and halo mass and the (mass-weighted) radial profiles of gas density, temperature, and pressure for all three simulations. The analysis includes comparisons to Planck tSZ observations and to the thermal pressure profile inferred from the Atacama Cosmology Telescope (ACT) measurements. We compare these tSZ data to simulations to interpret the measurements in terms of feedback and accretion processes in the CGM. We also identify as-yet unobserved potential signatures of these processes that may be visible in future measurements, which will have the capability of measuring tSZ signals to even lower masses. We also perform internal comparisons between runs with different physical assumptions. We conclude (1) there is strong evidence for the impact of feedback at  $R_{500}$ , but that this impact decreases by  $5R_{500}$ , and (2) the thermodynamic profiles of the CGM are highly dependent on the implemented model, such as cosmic-ray or AGN feedback prescriptions.

*Unified Astronomy Thesaurus concepts:* [Circumgalactic medium \(1879\)](#); [Sunyaev-Zeldovich effect \(1654\)](#); [Galaxy formation \(595\)](#); [Galaxy evolution \(594\)](#); [Astronomical simulations \(1857\)](#)

## 1. Introduction

With the development of observational tools and techniques, we are gaining a better understanding of the gas surrounding galaxies, known as the circumgalactic medium (CGM; see Tumlinson et al. 2017 for a recent review). The CGM and the intracluster medium (ICM) in galaxy clusters hold the majority of baryons as a result of multiple processes that distribute (and redistribute) them as diffuse gas within dark matter halos over cosmic time. The baryons in the CGM, both those suspended there due to accretion-shock heating and those expelled from galaxies and their interstellar medium (ISM), are believed to constitute the “missing baryons” (e.g., Persic & Salucci 1992; Fukugita et al. 1998; Bregman 2007). Therefore, the thermodynamic properties of the CGM afford useful information for understanding the physical processes governing galaxy formation and evolution.

Until recently, CGM studies have mostly relied on ultraviolet (UV) and optical absorption-line observations,

probing warm and cold components of the CGM (e.g., Suresh et al. 2017; Nelson et al. 2018a). The Sunyaev–Zel’dovich (SZ) effect (Sunyaev & Zeldovich 1970, 1972, 1980; see Mroczkowski et al. 2019a for a recent review) and soft X-ray emission have become powerful new tools to study the multiphase CGM’s hot gas component, whose temperature exceeds  $\sim 10^6$  K. These tools provide a calorimetric view of the ionized gas that complements absorption-line observations.

The thermal SZ (tSZ) effect is a spectral distortion of the cosmic microwave background (CMB) radiation caused by inverse Compton scattering of CMB photons by free electrons. The signal is proportional to the electron thermal pressure of the halo gas, integrated along the line of sight. Individual tSZ detections have been limited to massive clusters because of sensitivity and angular resolution, but stacked measurements of galaxy samples (e.g., Planck Collaboration et al. 2013a, *PCXI* hereafter; Greco et al. 2015) presented the mass-scaling relation down to halo masses below  $\sim 10^{13} M_{\odot}$ . Notably, the *PCXI* results suggested that the integrated tSZ flux within a radius  $R_{500}^{11}$  followed a self-similar dependence on halo mass over

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<sup>11</sup> The radius within which the mean overdensity is 500 times the critical density at a given redshift.

more than two orders of magnitude in mass. The mass scaling derived from a self-similar relation between the gas pressure profile and halo mass (Kaiser 1986, 1991) implied that the total thermal energy was governed primarily by the gravitational potential. This self-similarity suggested by PCXI appears counterintuitive because the nongravitational feedback mechanisms that regulate star formation and black hole growth in galaxies are expected to inject energy into and heat the CGM (e.g., Crain et al. 2009; Oppenheimer et al. 2010; Schaye et al. 2010). Specifically, stellar winds and supernovae (“stellar feedback” due to star formation) and active galactic nuclei (AGNs; “AGN feedback”) are expected to eject gas from the disks of galaxies into the CGM or out past it into the intergalactic medium (IGM) in a manner that depends on various halo properties such as halo mass, physics of star formation, and black hole mass accretion efficiency. However, the PCXI result was driven by the extrapolation of results measured at large radii ( $5R_{500}$ ) to  $R_{500}$ , assuming a fixed pressure profile model of Arnaud et al. (2010) based on the Nagai et al. (2007) profile. Also, Hill et al. (2018) pointed out that ignoring the signal from nearby halos (“two-halo term”) in the modeling could also have biased the inference toward a self-similar model. Therefore, a careful comparison between the tSZ observation and models, with an emphasis on the distribution of the hot gas and its properties, is required.

Observations using tSZ have already proven their worth in probing feedback (e.g., Verdier et al. 2016; Crichton et al. 2016; Spacek et al. 2018; Meinke et al. 2021). Battaglia et al. (2017) furthermore demonstrated that combined measurements of the tSZ and the kinetic SZ (kSZ) effects provide constraints on thermodynamic processes affecting the gas. Recently, Schaan et al. (2021) and Amodeo et al. (2021) cross-correlated the Atacama Cosmology Telescope (ACT) CMB survey maps (Naess et al. 2020) with the Baryon Oscillation Spectroscopic Survey (BOSS; Ahn et al. 2014) CMASS (“constant stellar mass”) galaxy catalog to obtain stacked tSZ and kSZ measurements. Comparing the radial thermodynamic profiles from the measurements to the predictions from cosmological simulations, they found that the simulations tend to underpredict the gas pressure and density, notably around the halos’ outer regions. Future millimeter-wave SZ instruments with improved angular resolution and sensitivity will provide data from halos down to galaxies and group scales (Mroczkowski et al. 2019b), enabling us to effectively probe feedback mechanisms in these lower-mass halos (Battaglia et al. 2019). Such instruments include next-generation CMB experiments (e.g., CMB-S4, Simons Observatory; Abazajian et al. 2016, 2019; Abitbol et al. 2017; Ade et al. 2019) and single large-aperture telescopes (e.g., CSST, CMB-HD, AtLAST; Golwala 2018; Klaassen et al. 2019; Sehgal et al. 2019).

While instrumentation and telescope development have enabled higher-sensitivity and higher-resolution observations, theoretical approaches have made significant progress as well. These include analytical studies of the tSZ and kSZ effects due to the halo gas that can constrain the baryon distribution around the halos (e.g., Natarajan & Sigurdsson 1999; Singh et al. 2015, 2016). On the other hand, current cosmological simulations have been successful in reproducing many observed galaxy properties, improving our understanding of galaxy formation and evolution (see Vogelsberger et al. 2020 for a recent review). However, these simulations often do not resolve the physical scales necessary to compute many

astrophysical processes and instead resort to “subgrid” prescriptions. The subgrid prescriptions are coarse-level approximations of what physical mechanisms are occurring on finer resolution scales than are self-consistently modeled in the simulation. While these subgrid models are calibrated against key observables, they are still a major source of uncertainty in computational models of galaxy evolution, so it behooves us to analyze multiple simulations with different subgrid prescriptions. These physically motivated recipes give predictions for CGM properties that vary significantly depending on the models (e.g., van de Voort & Schaye 2012; Hummels et al. 2013). In particular, a comparison of tSZ simulations to data can distinguish which prescriptions are correct and thereby inform our understanding of the CGM and feedback mechanisms that affect it.

In this work, we study the tSZ effect and relevant gas properties using state-of-the-art simulations. These include two large-volume simulations, Illustris-TNG (Nelson et al. 2018b; Pillepich et al. 2018a; Marinacci et al. 2018; Naiman et al. 2018; Springel et al. 2018, The Next Generation; TNG hereafter) and Evolution and Assembly of GaLaxies and their Environments (EAGLE; Schaye et al. 2015; McAlpine et al. 2016; The EAGLE team 2017), and one zoom-in simulation, Feedback In Realistic Environments-2 (FIRE-2; Hopkins et al. 2018; Orr et al. 2020). In Section 2, we describe the simulations and how we calculate the thermodynamic properties of the CGM and the tSZ signal. In Section 3, we compare the tSZ signal predicted by different models to the observations and perform internal comparisons between simulations. In particular, we (1) study scaling relations for integrated tSZ flux using TNG and EAGLE, separated by galaxy type, (2) compare EAGLE’s predictions to the radial profiles inferred recently from ACT data, and (3) discuss radial profiles of Milky Way-sized ( $\sim 10^{12} M_{\odot}$ ) halos at  $z=0$  from TNG, EAGLE, and FIRE-2. For the latter, we include the recent FIRE-2 simulations that model cosmic rays (CRs; Hopkins et al. 2020; Ji et al. 2020). We then (4) discuss the effect of AGN feedback on the tSZ flux by comparing EAGLE simulations with and without the AGN feedback prescription. Finally, we summarize the implications of our results in Section 4 and discuss possible synergies of tSZ observations with other CGM probes in Section 5.

## 2. Simulations and Methods

### 2.1. Thermal SZ Effect

The tSZ effect is a distortion of the thermal frequency spectrum of the CMB caused by inverse Compton scattering of CMB photons with free electrons in any hot gas (including the CGM). The net effect is a shift of photons from lower to higher frequencies. The resultant spectrum is unique and has the same frequency dependence for all objects in the limit that the electrons are nonrelativistic.<sup>12</sup> The amplitude of the distortion is proportional to the electron pressure integrated along the line

<sup>12</sup> Relativistic corrections cause the frequency dependence to depend on the electron temperature for temperatures approaching the electron rest energy ( $T \gtrsim 10^8$  K; e.g., Chluba et al. 2012), but the effect is negligible for typical CGM temperatures.

of sight, as characterized by the Compton- $y$  parameter,

$$\begin{aligned} y &\equiv \frac{\sigma_T}{m_e c^2} \int P_e dl \\ &= \frac{\sigma_T}{m_e c^2} \int n_e k_B T_e dl, \end{aligned} \quad (1)$$

where  $n_e$  is the electron number density,  $\sigma_T$  is the Thomson cross section,  $k_B$  is the Boltzmann constant, and  $T_e$  is the electron temperature;  $m_e$  is the electron rest mass and  $c$  is the speed of light. We describe the details of computing the tSZ signal using the simulation data set in Appendix A. The literature, including PCXI, often quotes the parameter integrated within a sphere of a radius  $R_{500}$ :

$$\begin{aligned} Y_{R_{500}} D_A^2(z) &\equiv \int_0^{R_{500}} n_e \sigma_T \frac{k_B T_e}{m_e c^2} dV \\ &= \int_0^{R_{500}} n_e \sigma_T \frac{k_B T_e}{m_e c^2} 4\pi r^2 dr, \end{aligned} \quad (2)$$

where  $D_A(z)$  is the angular diameter distance to the given redshift. In order to put sources at different redshifts on the same footing (the equivalent of converting from flux to luminosity for more traditional measurements), we scale to  $z=0$  and a fixed angular diameter distance of 500 Mpc,

$$\check{Y}_{R_{500}} = Y_{R_{500}} E^{-2/3}(z) [D_A(z)/500 \text{ Mpc}]^2, \quad (3)$$

where  $E(z)$  is the Hubble parameter  $H(z)$  normalized by  $H_0$ , which is approximated by  $E^2(z) = \Omega_m(1+z)^3 + \Omega_\Lambda$ . We express  $\check{Y}_{R_{500}}$  in units of  $\text{arcmin}^2$ .<sup>13</sup> Because the temperature is proportional to  $M^{2/3}$  for a virialized halo, the tSZ relation is dependent only on  $M$ , i.e.,  $\sim M^{5/3}$ , in the self-similar limit, in which gravity and thus mass determine all the halo parameters.

However, the spherically integrated  $Y_{R_{500}}$  is not a direct observable due to the instrumental beam size and the contribution of all gas along the line of sight. Because of Planck’s several-arcminute beam, PCXI actually measured the signal within a cylindrical volume of projected radius  $5 \times R_{500}$  and then extrapolated to  $Y_{R_{500}}$  assuming a universal electron pressure radial profile (Arnaud et al. 2010). In addition, the line-of-sight integration can be affected by the emission from nearby halos (“two-halo term”), thermal emission from dust in the host galaxies, the viewing angle, or other effects. The effect of the two-halo term on calculating the radial density and pressure profiles was discussed in several works (e.g., Vikram et al. 2017; Hill et al. 2018; Moser et al. 2021), and it becomes nonnegligible beyond a few times the virial radius. Following their approach, we include gas particles that do not belong to a central subhalo for the tSZ calculation in this work and discuss this issue in Appendix B.

## 2.2. Simulations

We use three simulation data sets, namely TNG, EAGLE, and FIRE-2, to compare tSZ signatures and gas radial profiles among galaxy samples. The simulations were tuned to reproduce various observed galaxy properties, such as the stellar-mass function and galaxy size, employing different subgrid prescriptions (Schaye et al. 2015; Pillepich et al.

2018b; Hopkins et al. 2018). Each of the simulations offers predictions for the distribution of the CGM in and around halos and for its thermodynamic properties that are highly dependent on its assumed model.

We mainly use the publicly available TNG100-2<sup>14</sup> (Nelson et al. 2019) and EAGLE<sup>15</sup> Ref-L100N1504 in similar comoving volumes ( $\sim 100^3 \text{ cMpc}^3$ ; cMpc: comoving Mpc). We focus our analysis on the present-day ( $z=0$ ) snapshots from both simulations. To test the impact of AGN feedback, we also use lower-resolution, smaller-volume EAGLE simulations with and without feedback (Ref-L0050N0752 and NoAGN-L0050N0752), which use a comoving volume of  $50^3 \text{ cMpc}^3$  with  $752^3$  smoothed particle hydrodynamics (SPH) and dark matter (DM) particles. TNG and EAGLE provide halo (group) and subhalo (galaxy) catalogs identified using the friends-of-friends (FoF; Huchra & Geller 1982; Davis et al. 1985) and subfind algorithms (Springel et al. 2001). We only identify central “subhalos,” which are the equivalent of central galaxies, as galaxies for the analysis so that we can be assured that each galaxy’s properties are driven by its own evolution rather than being dominated by a nearby larger galaxy. We do, however, include all nearby particles within a simulation volume to analyze gas properties, even if they are not bound to the specific halo/subhalo, because the CGM of a central galaxy and its satellites are not separate. We will describe the rationale for including these “two-halo” effects in detail in Appendix B.

We complement these recent state-of-the-art cosmological simulations with the “zoom-in” FIRE-2 simulation, which explores much smaller physical scales than the other two large-volume simulations and focuses on individual halos because its finer physical resolution has already provided an additional useful understanding of CGM properties (e.g., Stern et al. 2021). In particular, in Section 3.4, we compare the gas radial profiles of the Milky Way-sized galaxies in FIRE-2 to average profiles of galaxies of similar mass from TNG and EAGLE. We assess the impact of CRs by also comparing them to the recent FIRE-2 run with CR treatment (Ji et al. 2020).

CGM properties in these simulations have been analyzed in several prior works. Nelson et al. (2018a) analyzed TNG simulation data to study the distribution of high-ionization species O VI, O VII, and O VIII and their physical properties, which could be observed by UV and X-ray spectroscopy. Davies et al. (2019) studied the relation between the gas fraction and the halo mass of present-day  $\sim L^*$  galaxies in EAGLE. They showed that the central black hole mass and the halo gas fraction are strongly negatively correlated at a fixed halo mass and that the soft X-ray luminosity and the tSZ flux display similar correlations with the black hole mass and the star formation rate (SFR). Davies et al. (2020) extended this work to TNG, finding that the scatter in CGM mass fraction is strongly correlated with the specific SFR (sSFR) in both EAGLE and TNG while the CGM fractions in low-mass halos are considerably higher in TNG than in EAGLE. This result implies that tSZ observations of galaxy samples separated by their physical properties will give useful information about the halo gas and distinguish models.

van de Voort et al. (2016) studied the tSZ signal and the soft X-ray emission in the original FIRE (FIRE-1; Hopkins et al. 2014) simulation to understand the effect of stellar feedback.

<sup>13</sup> Some authors do not scale to  $D_A = 500$  Mpc, resulting in a quantity in  $\text{Mpc}^2$  rather than  $\text{arcmin}^2$ .

<sup>14</sup> <http://www.tng-project.org/>

<sup>15</sup> <http://icc.dur.ac.uk/Eagle/>

(The FIRE-1 simulation did not implement AGN feedback.) They found that the tSZ flux (integrated only out to  $R_{500}$  but including all the gas along the line of sight inside that radius in the simulation volume) of halos whose masses are below  $10^{13} M_{\odot}$  deviate from the PCXI self-similar scaling relation. They were able to explain this suppression in these low-mass objects by their reduced hot gas ( $T > 10^4$  K) fraction, which is also observed to be redshift independent, unlike the baryon and total gas fractions. They explained that the little dependence on the redshift was because high-redshift halos contained more cold gas than hot gas, and they lost a smaller fraction of baryons than the low-redshift halos, canceling out the effect on the hot gas fraction.

### 3. Thermal SZ Signal in Cosmological Simulations

The tSZ signal offers a way to probe hot gas components around the halos by measuring the integrated electron pressure. It has been observed with several millimeter instruments and wide-field, high-sensitivity CMB surveys. The observational effort includes experiments such as Planck (Planck Collaboration et al. 2011, 2013a), ACT (Henderson et al. 2016), and the South Pole Telescope (SPT; Benson et al. 2014) and will continue in the future with the Simons Observatory (SO; Ade et al. 2019) and CMB-S4 (Abazajian et al. 2016, 2019; Abitbol et al. 2017).

A key observational diagnostic is the scaling relation between the tSZ flux and halo mass. Using a matched multifilter (Melin et al. 2006), PCXI reported a nearly self-similar scaling relation for  $\sim 260,000$  locally brightest galaxies (LBGs) extracted from the Sloan Digital Sky Survey (SDSS; Abazajian et al. 2009). To convert observable stellar mass to halo mass, they generated a mock catalog using the Millennium simulation, selected LBG samples following the SDSS selection criteria, and derived the stellar-mass–halo-mass relation. The tSZ signal was measured down to a stellar mass of  $\sim 2 \times 10^{11} M_{\odot}$ , spanning a mass range of almost two orders of magnitude ( $M_{500}$  between  $\sim 10^{13}$  and  $\sim 10^{15} M_{\odot}$ ). Greco et al. (2015) later analyzed the data by applying aperture photometry instead of the matched filter, taking the stacking bias into account, and obtained results consistent with PCXI. Pratt & Bregman (2020) found a self-similar scaling relation for X-ray-selected groups and clusters with masses above  $10^{13.4} M_{\odot}$ .

This self-similarity might seem to contradict the intuition that nongravitational processes, such as stellar and AGN feedback, should cause a mass-dependent deviation from self-similarity with greater deviation at lower mass (e.g., Le Brun et al. 2015; van de Voort et al. 2016). We will see below that simulations can exhibit self-similarity or deviate from it depending on the radial scale over which quantities are defined. It is worth noting in this context that, because of the large instrumental beam, the Planck result actually measured the tSZ flux on scales out to  $5R_{500}$  and extrapolated back to  $R_{500}$  using the filter template based on the pressure profile from Arnaud et al. (2010).

In parallel with the observations, the tSZ effect has been studied with several hydrodynamic simulations (e.g., Le Brun et al. 2015; Barnes et al. 2017; McCarthy et al. 2017; Lim et al. 2018, 2021; Davies et al. 2019). For example, Le Brun et al. (2015) performed mock tSZ observation using the cosmological simulations. They studied several variations from the “reference” model that included metal-dependent radiative cooling and stellar feedback. Their “nonradiative”

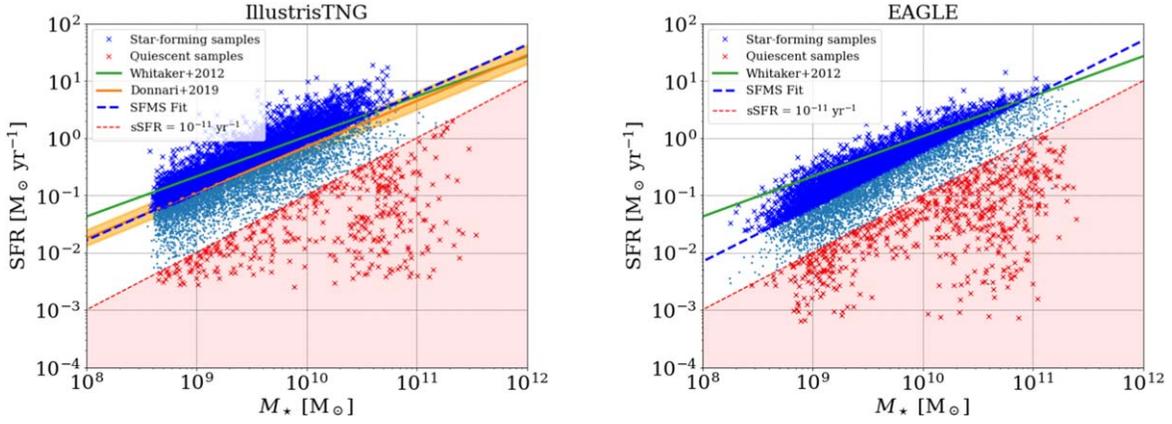
model only invoked the UV and X-ray background, without cooling and feedback prescriptions, while the AGN models included AGN feedback with varying temperature treatments. They found that both the tSZ signal and radial pressure profile change substantially among the models. An important implication is that extrapolation from  $5R_{500}$  to  $R_{500}$  based on a fixed  $\beta$  pressure profile, as in the Planck analysis, can significantly bias the results. Higher angular resolution (arcminute or better), or full forward modeling from simulation to observation, is therefore critical for probing the CGM in galaxy-scale halos.

Lim et al. (2021) compared the tSZ measurements from the Planck data to a number of hydrodynamical simulations. They analyzed the samples within the mass range of  $M_{500} \sim 10^{12-14.5} M_{\odot}$  and showed that the integrated tSZ flux  $\tilde{Y}_{R_{500}}$  of the models (Illustris, TNG-300, EAGLE, and Magneticum) starts to deviate from the self-similar relation below  $M_{500} \sim 10^{13} M_{\odot}$ . Their comparison showed that the simulated tSZ flux deviates to a greater extent as the halo mass  $M_{500}$  decreases, and each simulation gives a different amount of discrepancy from the self-similar relation due to the feedback model they adopted. These preliminary studies demonstrate the potential of SZ observations to add a new kind of constraint on galaxy formation models.

In this section, we analyze the tSZ signal of the halos in TNG, EAGLE, and FIRE-2 simulations. Because the large-volume cosmological simulations provide sufficient samples for a statistical study, we first explore the mass-scaling relation with TNG and EAGLE, especially by separating galaxy types. Then, we compare the radial profiles of the thermodynamic properties for  $\sim 10^{12} M_{\odot}$  galaxies from the three simulations, including the zoom-in FIRE-2 simulation. Equation (3) indicates that the tSZ flux  $\tilde{Y}_{R_{500}}$  is proportional to  $(\Omega_b/\Omega_m)h^{2/3}$  once the self-similar pressure profile is assumed (Arnaud et al. 2010; Lim et al. 2021). Because the PCXI analysis and the simulations we use in this analysis adopted different cosmological parameters (Komatsu et al. 2011; Planck Collaboration et al. 2014, 2016), we apply the correction to scale the measurements to the PCXI result.

#### 3.1. Sample Selection in TNG and EAGLE

We use simulated galaxies to compare the mass-scaling relation with the tSZ flux, depending on galaxy type. The previous observational and simulation studies of tSZ flux and halo mass imply the existence of a reliable self-similar relation down to  $\sim 10^{13} M_{\odot}$ . To see whether the self-similarity extends to lower masses, we limit our galaxy samples to the halo-mass range of  $10^{11.0} M_{\odot} < M_{500} < 10^{13.5} M_{\odot}$  in order to explore the effect of feedback in the simulations. (Note that the simulation study of Lim et al. 2021 used the mass range of  $M_{500} \sim 10^{12-14.5} M_{\odot}$ .) To understand how galaxy type affects the gas distribution around halos, and thus the tSZ signal, we categorize our galaxy samples from the TNG and EAGLE simulations according to their SFR. In general, we divide galaxies into star-forming and passive (quenched or quiescent galaxies with little SFR in the local universe) systems. Star-forming galaxies tend to reside in halos less massive than passive galaxies at a given stellar mass (Mandelbaum et al. 2016). Our analysis uncovers the dependence of the scaling relation on galaxy type, as well as the relations expected in the low-mass regime, where future tSZ observations will explore. Furthermore, the SFR is an indicator



**Figure 1.** Star-forming and quiescent galaxy samples used in our analysis for TNG100-2 (left) and EAGLE Ref-L100N1504 (right) at  $z = 0$ . We use well-resolved central galaxies with  $10^{11.0} M_{\odot} < M_{500} < 10^{13.5} M_{\odot}$ . We designate galaxies above the SFMS fit as star-forming (blue) and those with  $sSFR < 10^{-11} \text{ yr}^{-1}$  as quiescent (red). The galaxies below the SFMS fit but with  $sSFR > 10^{-11} \text{ yr}^{-1}$  are not taken into account to better contrast the effect of galaxy type. The green line shows a fit to the observations (Whitaker et al. 2012). The orange line is the SFMS fit to TNG-300, presented in Donnari et al. (2019) with  $\alpha(0) = 0.80 \pm 0.01$  and  $\beta(0) = -8.15 \pm 0.11$ .

of stellar feedback assuming that the energy from galactic winds is proportional to the instantaneous SFR (e.g., Pillepich et al. 2018b; Weinberger et al. 2018). Our analysis tests the expectation that the SFR affects the tSZ signal close to the host galaxy.

The relation between stellar mass and SFR has been studied for both TNG and EAGLE (Furlong et al. 2015; Donnari et al. 2019). Star-forming galaxies follow a relation in the stellar-mass ( $M_*$ )–SFR plane (e.g., Karim et al. 2011) known as the star-forming main sequence (SFMS). Donnari et al. (2019) identified the SFMS in the  $M_*$ –SFR plane for the TNG simulations. They fit the SFR of the star-forming samples as a function of the stellar mass,

$$\log\left(\frac{\langle SFR \rangle_{\text{sf-ing galaxies}}}{M_{\odot} \text{ yr}^{-1}}\right) = \alpha(z) \log\left(\frac{M_*}{M_{\odot}}\right) + \beta(z), \quad (4)$$

where they defined the SFR of a halo as an integration of individual SFR within twice the spherical stellar half-mass radius and the stellar mass as a sum of all the stellar particles within the same radius. To make a coherent comparison, we use the SFR and the  $M_*$  measured within a 30 (physical) kpc aperture for both TNG and EAGLE, following the common EAGLE analysis (e.g., Davies et al. 2020).

We divide galaxies from TNG and EAGLE into star-forming and passive samples. First, we identify the SFMS for central galaxies, whereas Donnari et al. (2019) did not distinguish centrals from satellites. We consider only systems with more than 50 gas particles and stellar particles each to ensure that the galaxies are well resolved. Then, we classify as quiescent those galaxies with  $sSFR$  below  $10^{-11} \text{ yr}^{-1}$  (e.g., Wetzell et al. 2013), where  $sSFR = SFR/M_*$ . We fit the SFMS using the rest of the galaxies, performing a linear fit to the median SFRs of  $M_*$  in the range between  $10^9$  and  $10^{10}$ , separated by 0.2 dex bins. We define the galaxies above the MS fit as star-forming samples in this analysis to better contrast the results depending on galaxy types.

Figure 1 shows the  $M_*$ –SFR plane populated with the present-day central galaxies in TNG (snapshot 99) and EAGLE (snapshot 28) and our aforementioned samples. The criteria we use give 4178 star-forming and 1590 quiescent galaxies for TNG, and 3355 star-forming and 907 quiescent galaxies for

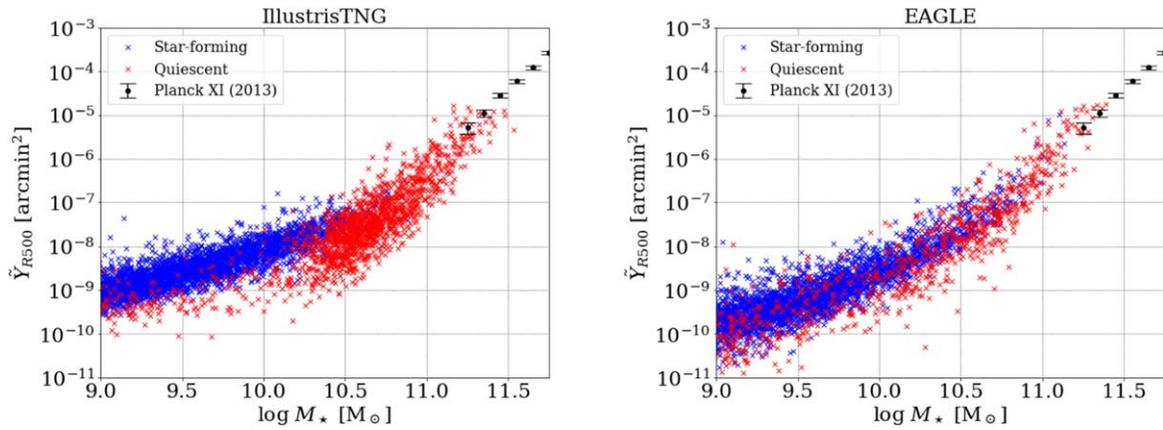
EAGLE. Because TNG has a  $\sim 30\%$  larger volume than EAGLE, it accordingly contains more galaxies. However, these are of similar order and will not affect the interpretation

### 3.2. Integrated tSZ Signal: Mass-scaling Relation

In this section, we examine the tSZ mass-scaling relation ( $Y$ – $M$  relation) for the TNG and EAGLE galaxy samples, divided by SFR. The aims of showing the  $Y$ – $M$  relation using the simulated galaxy samples are to (1) compare results to the PCXI data, where the relation shows little departure from a single power law with a slope of  $5/3$  and (2) investigate how the relation depends on the type of galaxy (star-forming versus quiescent) and understand how the feedback physics included in each simulation can affect the relation.

We spherically integrate the tSZ signal out to  $R_{500}$  to calculate  $Y_{R_{500}}$  for each of the galaxy samples. The  $Y$ – $M$  relation is often presented as a function of the halo mass  $M_{500}$  to relate the observable tSZ flux to the dark matter halo properties. However, stellar mass is the more directly measurable property, and observations determine the tSZ to stellar-mass relation. The stellar-mass–halo-mass conversion is highly dependent on galaxy properties, especially on color. Figure 2 shows the spherically integrated  $Y_{R_{500}}$  as a function of  $M_*$ , with the PCXI measurements overlaid. The simulations and the PCXI observation overlap around  $10^{11.25} M_{\odot} < M_* < 10^{11.5} M_{\odot}$ . Note that the PCXI points in their work flatten out below  $M_* \sim 10^{11} M_{\odot}$  because of dust contamination. The  $M_* = 10^{11.25} M_{\odot}$  bin of the Planck measurement was made at  $3.5\sigma$  detection, and the lower-mass bins show weaker detection significance; therefore, we only include the points above  $M_* = 10^{11.25} M_{\odot}$  in Figure 3.

Figure 3 shows the simulated tSZ flux as a function of halo mass  $M_{500}$  and compares it to the PCXI scaling relation. As described earlier, PCXI measured the tSZ signal out to  $5R_{500}$  and converted it to  $Y_{R_{500}}$  assuming the universal pressure profile (UPP; Equation (22) in Arnaud et al. 2010). This rescaling of the tSZ flux to go from  $Y_{R_{500}}$  to  $Y_{5R_{500}}$  corresponds to a factor of 1.796. We present the scaling relation in terms of both  $Y_{R_{500}}$  and  $Y_{5R_{500}}$ . The simulated galaxies lie below the PCXI scaling relation for  $Y_{R_{500}} - M_{500}$ , and the discrepancy grows as the mass decreases for both types of galaxies and for both simulations. If, on the other hand, we consider the scaling relation in  $Y_{5R_{500}} - M_{500}$ , we find good agreement.



**Figure 2.** tSZ flux as a function of stellar mass for TNG (left) and EAGLE (right). Blue and red crosses indicate the star-forming and quiescent samples as defined in Figure 1. Black circles show the PCXI measurements, and the error bars indicate the statistical error. Note that our samples do not cover the entire range of the stellar mass from the PCXI observation because we limited our galaxy mass range to  $10^{11.0} M_{\odot} < M_{500} < 10^{13.5} M_{\odot}$ .

To test the explanation of this “aperture size” effect in the tSZ flux measurement, we show the ratio  $Y_{5R_{500}}/Y_{R_{500}}$  as a function of halo mass and separated by galaxy type in Figure 4. The stark difference between these two scaling relations ( $Y_{R_{500}} - M_{500}$  and  $Y_{5R_{500}} - M_{500}$ ) could be explained qualitatively by feedback driving gas from small radii ( $\sim R_{500}$ ) out to larger radii such that all the gas and energy one expects are contained inside  $5R_{500}$  but not inside  $R_{500}$ . We observe two important trends. The first is that, for all galaxies, the ratio converges to the universal pressure profile value as the halo mass increases. The second is that the discrepancy with the UPP is larger for quiescent galaxies. Both of these trends support the feedback explanation. The trend with mass suggests that, as the halo mass grows, gravity wins out over feedback effects driving gas outward, resulting in better agreement with the UPP (derived from massive clusters; see below), or, equivalently, feedback is increasingly effective in overcoming gravity at  $R_{500}$  as the halo mass decreases. The observation of a larger discrepancy for quiescent galaxies is consistent with the idea that feedback has had its full effect in that class, blowing the bulk of the gas fuel for star formation out into the far CGM, while star-forming galaxies are still somewhere along that evolutionary trend, with less of the gas blown out past  $R_{500}$ , making it possible for them to continue forming stars at a high rate.

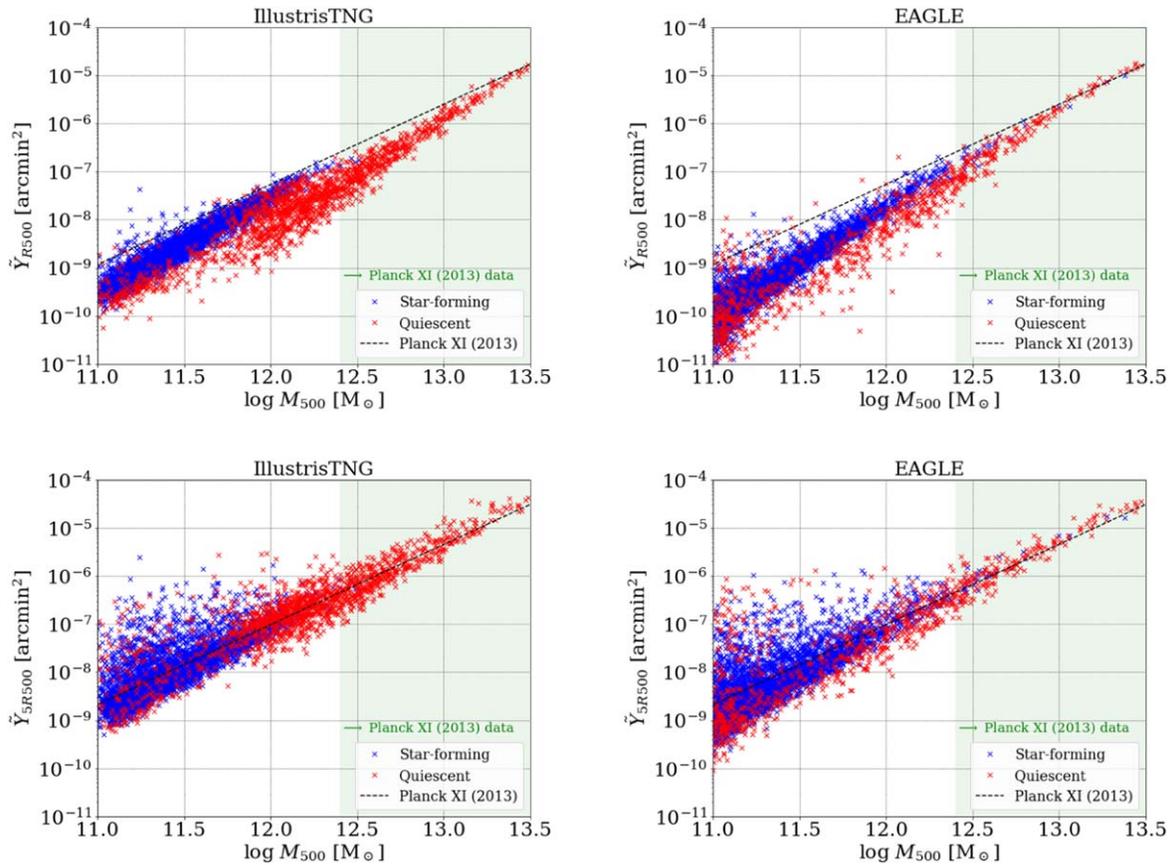
These deviations from self-similarity at  $R_{500}$  occur simply because the UPP, on which the extrapolation from  $Y_{5R_{500}}$  to  $Y_{R_{500}}$  is based, is not correct for galaxies. One must remember that it was obtained from a sample of massive clusters, in which feedback plays a much smaller role (fractionally). In the mass range for which they overlap,  $10^{11.5} < M_{500} < 10^{12.5}$ , the relative locations of the star-forming and quiescent galaxies in Figures 3 and 4 are also illuminating. The much larger  $Y_{5R_{500}}/Y_{R_{500}}$  ratio for quiescent galaxies suggests that a much larger fraction of their gas has been heated and pushed out to a larger radius than for star-forming galaxies. This is consistent with the idea that star formation in quiescent galaxies was quenched by feedback from star formation and AGN activity that caused this heating/pushing of the gas (though the details of the model, including the timescale it takes to set up this inequality in the profile, need to be further studied, and it is outside the scope of this work). Additionally, the growth of the ratio and the difference between star-forming and quiescent galaxies at smaller masses indicate that the effect is fractionally larger at lower masses. While the details are quite dependent on

the simulation, the effect is clearly present in both simulations and is consistent with expectations from the basic picture of feedback quenching star-forming galaxies by driving their gas out to a large radius and heating it.

### 3.3. Comparison with the ACT Stacked Profile

Recently, the ACT collaboration presented stacked tSZ and kSZ measurements by cross-correlating the ACT DR5 data set and the BOSS CMASS galaxy catalog (Schaan et al. 2021; Amodeo et al. 2021). From the SZ measurements, they inferred the gas-mass density and temperature, hence thermal pressure, profiles by fitting a parameterized generalized Navarro–Frenk–White (GNFW) profile (Battaglia et al. 2012). They compared the results to the NFW profile (Navarro et al. 1997), hydrodynamical simulations by Battaglia et al. (2010), and TNG galaxy samples as shown in Figure 5 (Figure 6 in Amodeo et al. 2021). Although the simulations matched the profiles deprojected from the observations well at radii smaller than  $\sim 2R_{200}$  ( $\sim 1$  Mpc), the discrepancy was not insignificant at larger radii, especially for the thermal pressure profile, thus temperature. They claimed the simulations underpredicted the gas density and thermal pressure and explained that it was due to the CGM gas in the outer region being less heated by the subgrid models the simulations employed. In this section, we present the stacked radial profiles using the EAGLE galaxies and compare those to the ACT best-fit results and the TNG profiles in Amodeo et al. (2021).

The CMASS galaxies are distributed within the redshift range  $0.4 < z < 0.7$ , and the galaxy samples used in the ACT analysis have median redshift  $z = 0.55$  and mean stellar mass of  $3 \times 10^{11} M_{\odot}$  (from Maraston et al. 2013 stellar-mass estimates), which correspond to a halo mass of  $M_{200} \sim 3 \times 10^{13} M_{\odot}$  using the Kravtsov et al. (2018) stellar-to-halo-mass conversion. Note that Amodeo et al. (2021) used  $R_{200}$  and  $M_{200}$  as virial radius and mass, whereas we use  $R_{500}$  and  $M_{500}$  throughout this paper. For their  $M_{200} \sim 3 \times 10^{13} M_{\odot}$  halo,  $R_{200}$  and  $R_{500}$  correspond to  $\sim 500$  kpc and  $\sim 350$  kpc, respectively. They selected red galaxies in the TNG simulation and calculated the average profile by weighting the samples using the stellar and halo-mass distributions of the observed galaxies (“TNG S” and “TNG H” in Figure 5). For example, they used nine log-spaced bins between  $10^{11.53}$  and  $10^{13.98} M_{\odot}$  for the halo-mass distribution. The mass catalog can be



**Figure 3.** tSZ flux as a function of halo mass  $M_{500}$ , integrated out to  $R_{500}$  (upper), and to  $5R_{500}$  (lower) in TNG (left column) and EAGLE (right column). The shaded green region shows the mass range of the galaxy samples used in the PCXI analysis, and their best-fit scaling relation is shown as the dotted line. The relation in the lower panels is rescaled by multiplying the original relation in the upper panels by 1.796 (ratio from the universal pressure profile).

downloaded as part of the MOP-C GT (“Model-to-observable projection code for Galaxy Thermodynamics”) package. We use the identical color cut ( $g - r \geq 0.6$ )<sup>16</sup> for the EAGLE central galaxies, using snapshots 22 ( $z = 0.62$ ) and 23 ( $z = 0.50$ ) to match the median redshift. We calculate the weighted average of the gas-mass and thermal pressure profiles using the halo-mass ( $M_{200}$ ) probability density function of Amodeo et al. (2021). Because the gas-mass density and pressure are proportional to  $M$  and  $M^{5/3}$ , respectively, it is necessary to use the correct mass distribution to estimate the stacked profiles (Moser et al. 2021). As described earlier, we include the two-halo term in calculating the radial profiles by taking the nearby particles within a volume into account.

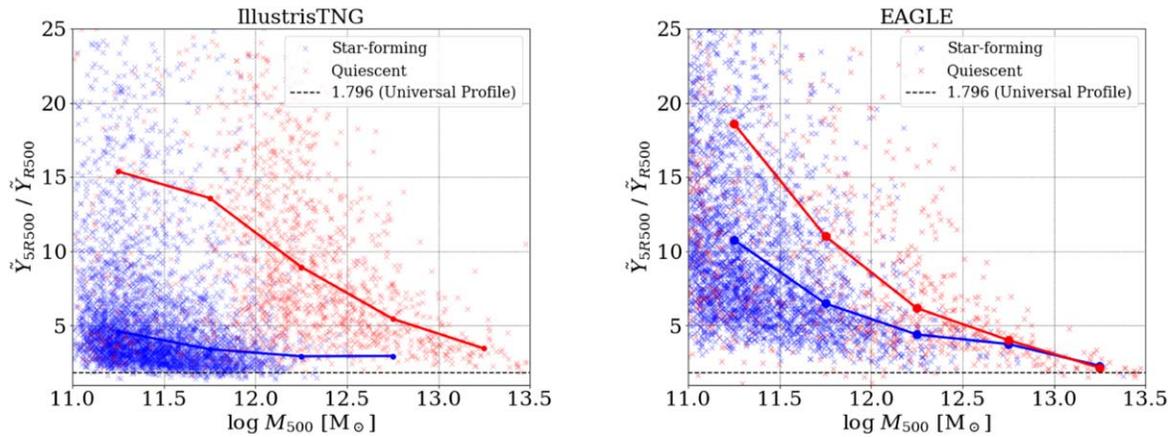
Figure 5 shows the gas-mass density and thermal pressure profiles inferred from the EAGLE data, compared to the ACT best-fit profiles as well as the TNG results in Amodeo et al. (2021). The EAGLE density and pressure profiles are nearly consistent with the TNG profiles. Both the EAGLE and TNG density profiles are within  $2\sigma$  of the ACT profile at most radial distances. The thermal pressure profiles show somewhat similar discrepancies among ACT, TNG, and EAGLE.

We could interpret the comparison of the simulation data (EAGLE and TNG) to the ACT profiles in two ways. First, the EAGLE profiles are approximately consistent with the TNG profiles. This supports the discussion in Amodeo et al. (2021) that the subgrid models in the cosmological simulations underpredict the gas density and pressure at large radii due to

insufficient heating. However, as we will see in Section 3.4, the FIRE-2 simulation with the CR treatment predicts even lower thermal pressure than TNG and EAGLE in general, primarily due to the temperature rather than the gas-mass density, and the most recent FIRE-2 subgrid model will not likely resolve this issue. Another possible interpretation is that feedback in TNG and EAGLE is too effective, that they blow too much gas out to a large radius. We argued earlier that feedback in TNG and EAGLE was not as effective in star-forming galaxies as in quiescent galaxies (Section 3.2). As we will see in Figure 6, the density and pressure for star-forming galaxies are higher than for quiescent galaxies in both simulations. This implies that perhaps the low gas-mass density and pressure profiles of TNG and EAGLE are results of very efficient feedback mechanisms.

Also, we could relate the comparison to the integrated tSZ flux and mass-scaling relation in Section 3.2. Figure 3 shows that the simulated tSZ flux integrated out to  $5R_{500}$  more or less agrees with the PCXI fit around the mean halo mass of the galaxy samples ACT analysis used ( $M_{200} \sim 3 \times 10^{13} M_{\odot}$  or  $M_{500} \sim 2 \times 10^{13} M_{\odot}$ ). However, if we use the ACT pressure profile in Figure 5 to calculate the tSZ flux out to  $5R_{500}$ , it is a few times higher than the PCXI tSZ flux measurement. The discrepancy is dependent on the two-halo term and could be due to the sample selection. For example, PCXI explicitly selected “LBGs” to construct a sample of central galaxies, which is defined as galaxies that do not have a brighter neighbor within 1 Mpc. An investigation into this effect would require fully modeling the selection functions of the PCXI and CMASS samples, which is beyond the scope of this work.

<sup>16</sup>  $g$  and  $r$  are the Sloan Digital Sky Survey (SDSS) magnitudes.



**Figure 4.**  $Y_{5R500}/\tilde{Y}_{R500}$  of the TNG and EAGLE galaxies. The dashed line shows the conversion factor 1.796 corresponding to the universal pressure profile. The red and blue dots connected by the solid lines indicate the median value of the ratios in each  $M_{500}$  mass bin. The mass bins are separated by 0.5 dex between  $10^{10.5} M_{\odot}$  and  $10^{13.5} M_{\odot}$ . Higher ratios of  $Y_{5R500}/\tilde{Y}_{R500}$  indicate that the hot gas has been pushed out to the outskirts of the galaxy halos by feedback.

### 3.4. Gas Radial Profiles of Milky Way–sized Galaxies

The study of integral quantities in Section 3.2 strongly suggests that there is a great deal of information about the impact of feedback on the CGM in the tSZ radial profile. Future SZ observations will be able to characterize the tSZ signal from Milky Way–sized or even less massive halos through stacking if individual detections are not possible. The CGM properties of Milky Way–mass halos have been studied with cosmological simulations. For example, Kauffmann et al. (2016, 2019) selected Milky Way–mass galaxies at  $z=0$  in Illustris and TNG simulations, based on both halo and stellar masses and explored the CGM structure, including the radial distribution of the gas. In this section, we explore the radial profiles of TNG and EAGLE galaxies in the mass range  $10^{11.75} < M_{500} < 10^{12.25} M_{\odot}$ , separated by galaxy type, and we also include FIRE-2 simulations of individual Milky Way–sized galaxies (see Table 1).

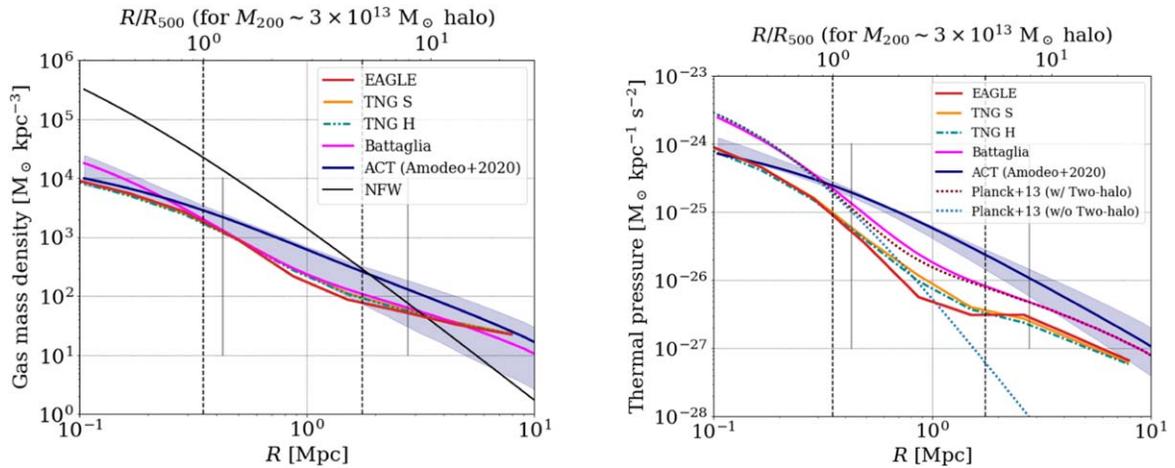
Figure 6 shows the mass-weighted radial profiles, including cumulative tSZ signal, gas-mass density, temperature, and thermal pressure. We selected the star-forming and quiescent samples in TNG and EAGLE based on the sample selection in Section 3.1 and limited the halo-mass ( $M_{500}$ ) range to  $10^{11.75} - 10^{12.25} M_{\odot}$ . This selection leaves 602 star-forming and 699 quiescent galaxies for the TNG and 392 star-forming and 226 quiescent galaxies for the EAGLE data set. Each of the dashed lines in the background shows an individual galaxy radial profile, and the solid thick curves display the median profiles. When calculating the radial profiles, we include all the gas particles around the center of each halo. However, a distinction between ISM and CGM gas particles can be made by using the SFR of the fluid elements (Davies et al. 2019, 2020). Because we mainly focus on the thermodynamic profiles around and beyond the virial radius, the separation of the ISM particles has a negligible effect on our results.

As shown in Section 3.2, star-forming galaxies tend to exhibit higher tSZ fluxes than quiescent galaxies when integrated out to  $R_{500}$ , although the difference is smaller or even negligible at  $5R_{500}$  in both TNG and EAGLE (Figure 3). This is consistent with the tSZ flux inferred for the  $10^{12} M_{\odot}$  halo at  $R=5R_{500}$  from the PCXI measurement (Figure 6). Again, the radial profiles reflect the intuition that the heating and expulsion of gas far out into the CGM has been much more extensive for quiescent galaxies than for star-forming ones. The

gas density is larger, and the temperature is lower close in ( $\log(R/R_{500}) < -0.5$ ) for star-forming galaxies compared to quiescent galaxies, which supply a reservoir of fresh fuel for star formation. Of course, because the thermal pressure (thus electron pressure) is proportional to the product of the gas density and temperature, these two effects can partially or completely cancel each other. At radii smaller than  $\sim 2R_{500}$  in TNG and at all radii in EAGLE, the gas density wins out over the temperature, yielding higher pressure for star-forming galaxies. At radii larger than  $\sim 2R_{500}$  in TNG, the gas density of the quiescent galaxies surpasses that of star-forming galaxies. This leads to higher thermal pressure for  $R \gtrsim 2R_{500}$ , and the tSZ fluxes of the quiescent and star-forming samples become comparable at  $3R_{500}$ .

The zoom-in FIRE-2 simulation (Hopkins et al. 2018) uses better spatial and mass resolutions than the TNG and EAGLE data sets, focusing on the physical processes of individual halos. van de Voort et al. (2016) used the FIRE-1 simulation, which only included the stellar feedback prescription without AGN feedback to study the tSZ effect and the X-ray emission properties of the halos. Incorporating the galaxy simulations’ redshift-dependent snapshots, they discovered that the halo’s hot gas fraction is dependent on the halo mass, with little redshift dependence. They pointed out that the tSZ measurements could observe this, with increasing suppression of the tSZ signal compared to the PCXI self-similar relation as the halo mass decreases below  $10^{13} M_{\odot}$  as a result of stellar feedback. Recently, the FIRE project presented the FIRE-2 simulations with a CR treatment (Hopkins et al. 2020; Ji et al. 2020). Ji et al. (2020) showed that the nonthermal CR pressure could dominate over thermal pressure in the CGM, supporting the gas, especially for Milky Way–sized halos. In these simulations, gas cooling is more effective than CR heating, and the gas temperatures are expected to decrease considerably compared to the reference simulations, while the gas-mass density profiles are not strongly affected. Because the tSZ directly measures the integrated thermal pressure, comparing the FIRE-2 reference and CR runs illustrates how the tSZ flux changes with the different subgrid models.

The radial profiles of the  $z=0$  (snapshot number 600) reference and CR simulation models are shown in the right column of Figure 6. We first observe that the reference simulations and CR runs show substantially different radial



**Figure 5.** Deprojected radial gas-mass density (left) and thermal pressure (right) profiles from the ACT and CMASS cross-correlation (Amodeo et al. 2021; blue line and band) compared with the TNG (orange and green), EAGLE (red), and Battaglia et al. (2010) simulation (magenta), and the NFW profile (Navarro et al. 1997, black). The dotted line in the thermal pressure plot shows the best-fit GFW pressure profile derived from the galaxy cluster observations in Planck Collaboration et al. (2013b). We show both the original pressure profile without a two-halo term (dotted blue) and with a two-halo term (dotted maroon) presented by Amodeo et al. (2021). The upper axis shows the radial distance in units of  $R_{500} = 350$  kpc, for an  $M_{200} \sim 3 \times 10^{13} M_{\odot}$  halo. The dashed vertical lines show  $R = R_{500}$  and  $R = 5R_{500}$ . The vertical gray bars indicate the radial ranges where the ACT SZ measurements were made (Schaan et al. 2021).

profiles, except for the gas-mass density profile, as already discussed in Ji et al. (2020). This implies that these CR-driven differences should be considered meaningful. CRs cause a substantial reduction in pressure and temperature out to  $3R_{500}$  because they transport energy and momentum created by star formation out to  $3R_{500}$  before slowing down and depositing their energy. We also see that the CRs greatly suppress the thermal pressure due to the decreased temperature, resulting in much lower tSZ flux than the reference runs at  $R_{500}$ .

When we compare the FIRE-2 to TNG and EAGLE, the FIRE-2 reference simulation’s radial profiles in all four quantities ( $Y$ , gas-mass density, pressure, and temperature) are generally in qualitative agreement with TNG and EAGLE. However, we cannot directly compare the FIRE-2 radial profiles to TNG and EAGLE at radii beyond  $\sim 1.5R_{500}$  because the zoom-in FIRE-2 simulations do not include the two-halo term from the gas particles outside their simulation volume. We observe that the tSZ radial profiles of the reference simulations flatten for  $R \gtrsim 2R_{500}$ . These tSZ profiles are similar to the TNG and EAGLE tSZ profiles when the two-halo term is not considered (Figure B1).

### 3.5. Effect of AGN Feedback in the EAGLE Simulation

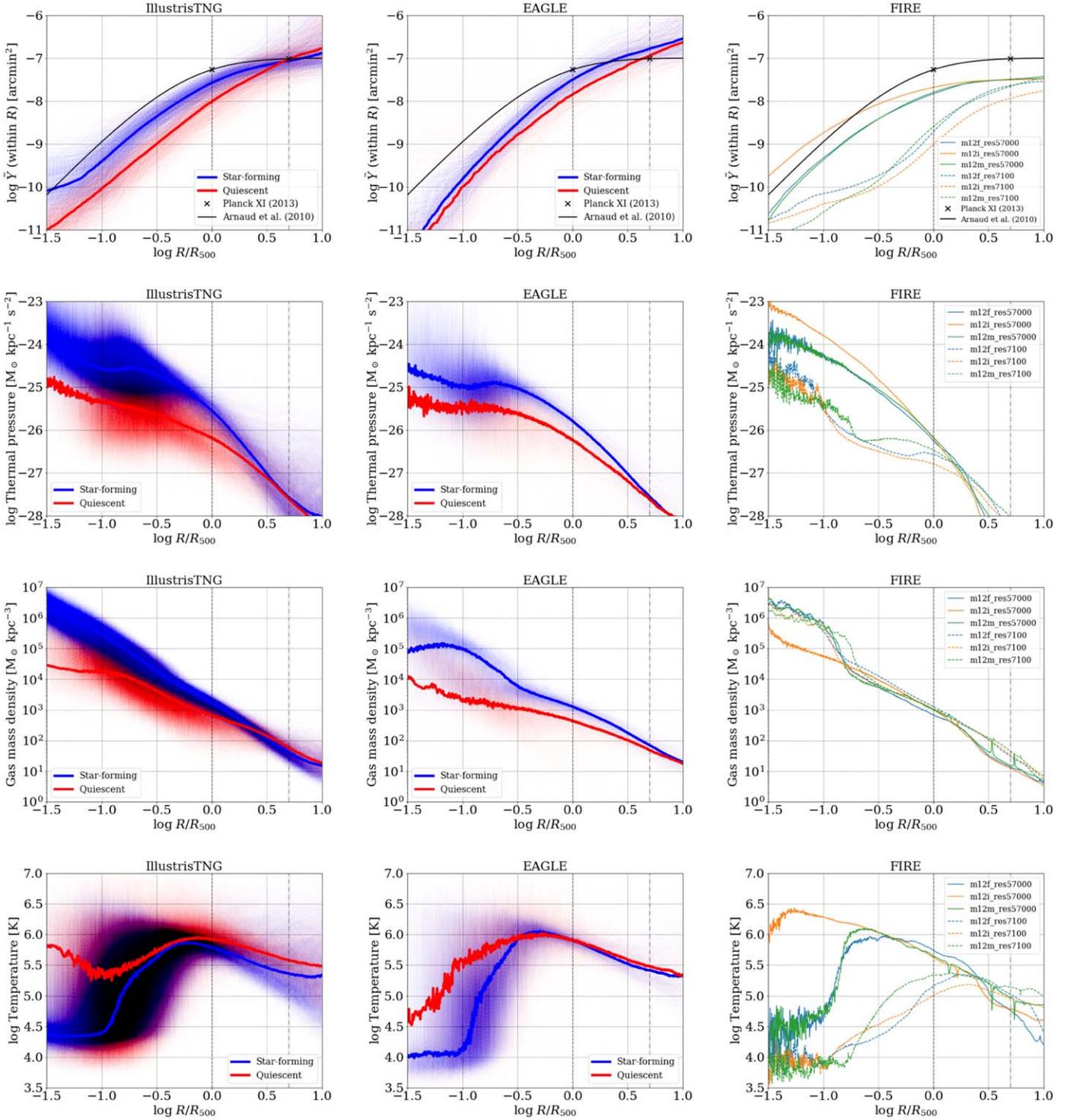
We compare the EAGLE simulations Ref-L0050N0752 (“reference”) and NoAGN-L0050N0752 (“NoAGN” simulation) at  $z = 0$  to explore the effect of the AGN feedback. They are in smaller  $50^3 \text{cMpc}^3$  volumes and use the same particle resolution as Ref-L100N1504. The NoAGN simulation does not include the subgrid model describing the AGN feedback. Davies et al. (2019) already compared the simulations by looking at the present-day halo baryon fraction, which is a sum of the gas and stellar mass over the total mass within the radius  $R_{200}$ . As expected, halos more massive than  $\sim 10^{12} M_{\odot}$  have higher baryon fractions in the NoAGN simulation because the stellar feedback cannot efficiently push out the gas in massive halos.

In Figure 7, we compare the tSZ flux measured within  $R_{500}$  and  $5R_{500}$ . We do not separate the galaxy types here because we are only interested in the behavior depending on the AGN

feedback prescription. As in Figure 3, the Compton- $y$  parameters integrated within the radius  $5R_{500}$  better reproduce the scaling relation of PCXI. The tSZ fluxes within  $R_{500}$  of the reference simulation seem to be lower than in the NoAGN run, in particular in the midmass range between  $10^{12}$  and  $10^{12.5} M_{\odot}$  (Figure 7, left panel). This becomes more evident in Figure 8, where we show the ratio of the tSZ fluxes within  $R_{500}$  and  $5R_{500}$ . The solid blue and orange lines indicate the median ratio of the galaxy samples within each mass bin, separated by 0.75 dex between  $10^{11.0}$  and  $10^{13.25} M_{\odot}$ . For both of the simulations, the ratios approach the value from the UPP as the halo mass increases. The median tSZ flux ratios of the reference run are greater than the NoAGN run, indicating that the hot gas has been pushed out beyond  $R_{500}$  by the AGN feedback. The median ratios of the NoAGN simulation samples between  $10^{12.5}$  and  $10^{13.25} M_{\odot}$  are nearly consistent with the self-similar value. As the halo mass decreases, the stellar feedback becomes efficient in both simulations, and the ratios deviate rapidly from the self-similar value.

In Figure 9, we present the average radial thermodynamic profiles of the  $10^{11.0} M_{\odot} < M_{500} < 10^{11.75} M_{\odot}$ ,  $10^{11.75} M_{\odot} < M_{500} < 10^{12.5} M_{\odot}$ , and  $10^{12.5} M_{\odot} < M_{500} < 10^{13.25} M_{\odot}$  samples in both of the EAGLE simulations shown above. First of all, we see that the tSZ flux of the two simulations at both  $R_{500}$  and  $5R_{500}$  is reduced compared to the PCXI interpolation as the halo mass decreases, although the flux at  $5R_{500}$  better agrees with the UPP profile than at  $R_{500}$ . As the halo mass increases, the tSZ radial profiles of the reference and NoAGN runs display a larger discrepancy, especially at  $R \lesssim R_{500}$ , due to the enhanced AGN feedback. At the highest mass bin, the tSZ profile of the NoAGN simulation is nearly consistent with the UPP profile, implying that gravity is able to overcome feedback and retain the gas close in without AGN feedback.

Comparing the gas-mass density and the temperature profiles in the middle and the right columns, we observe that the reduction in the flux at  $R \lesssim R_{500}$  for the reference simulation comes from the reduced gas-mass density (temperature is comparable or even higher), which is a result of the feedback processes pushing the gas out to larger radii. In the lowest-mass



**Figure 6.** Mean radial profiles of the TNG (left column) and EAGLE (middle column) galaxies separated by SFR in the mass range  $10^{11.75} < M_{500} < 10^{12.25} M_{\odot}$ . Blue and red curves show median radial profiles of the star-forming and quiescent samples, respectively. We show the cumulative tSZ flux, thermal pressure, gas-mass density, and temperature from top to bottom. The faint dotted lines are the radial profiles of the individual galaxies. The black crosses in the tSZ radial profile plots (top row) denote the tSZ signal for a  $10^{12} M_{\odot}$  mass halo at  $R_{500}$  ( $\sim 160$  kpc) and  $5R_{500}$ , inferred from the self-similar relation of PCXI, whose measurements are based on the halos at least an order of magnitude more massive than the Milky Way-sized halo. The solid black lines in those plots show the spherically integrated tSZ flux as a function of radius, derived from the universal pressure profile (Arnaud et al. 2010), normalized by the PCXI measurement. Radial profiles of the FIRE-2 galaxies (right column): The solid and dashed lines show the reference models and the CR included simulations, respectively. The FIRE-2 m12i, m12m, and m12f are all “star-forming galaxies” as opposed to quiescent.

bin, all the radial profiles of the NoAGN and reference simulations are comparable because stellar feedback dominates over AGN feedback in this mass range.

#### 4. Conclusions

We studied the thermodynamic properties of the CGM using the tSZ signal in TNG, EAGLE, and FIRE-2 simulations. We

**Table 1**  
FIRE-2 Simulation Summary

Name	$M_{\text{halo}}^{\text{vir}} (M_{\odot})$	$M_{*} (M_{\odot})$	$R_{\text{vir}} \text{ (kpc)}$	$M_{500} (M_{\odot})$	$R_{500} \text{ (kpc)}$
m12f	$1.6 \times 10^{12}$	$8.0 \times 10^{10}$	306	$1.2 \times 10^{12}$	163
m12i	$1.2 \times 10^{12}$	$6.5 \times 10^{10}$	275	$9.1 \times 10^{11}$	150
m12m	$1.5 \times 10^{12}$	$1.2 \times 10^{11}$	301	$1.1 \times 10^{12}$	159

**Notes.** List of the FIRE-2 zoom-in simulations used in this work. The physical quantities are measured in the reference simulation data set (see Hopkins et al. 2018 for details). We used both the reference and CR runs for each of the models.

find that the level of agreement of simulations over the mass range  $M_{500} = 10^{11}$  to  $3 \times 10^{13} M_{\odot}$  with a self-similar fit inferred from Planck data down to  $M_{500} \sim 10^{13} M_{\odot}$  ( $M_{*} \sim 10^{11} M_{\odot}$ ) is dependent on the aperture size used for the tSZ measurement. We interpret this dependence as evidence that feedback has a significant impact on CGM thermodynamic properties at  $R_{500}$  (significant deviations from self-similarity) but that these effects have subsided by  $5R_{500}$  (self-similarity is recovered). We also find that the effect is more pronounced in quiescent than in star-forming galaxies, suggesting that the impact of feedback at  $R_{500}$  is still in process for star-forming galaxies while it is complete for quiescent ones.

We find that agreement on radial profiles of gas density and thermal pressure between simulations and ACT analysis at  $M_{200} \sim 3 \times 10^{13} M_{\odot}$  ( $z \sim 0.55$ ) is also radius dependent. The disagreement at large radius is consistent with prior suggestions that subgrid heating in the simulation at large radius is insufficient. Still, it could also be a sign that feedback in the simulations is too effective at these radii.

We also compare the simulated radial profiles of gas density, temperature, and thermal pressure for Milky way–mass galaxies by separating star-forming and quiescent galaxies. We find support from all three simulations for the above interpretation. We also find from FIRE-2 simulations that the inclusion of CRs further enhances these effects, presumably because of the nonthermal pressure supplied by the CR.

Lastly, we compare EAGLE simulations with and without AGN feedback. We find that the impact of AGNs on the  $Y$ – $M_{500}$  scaling relation and on the radial profiles is most significant near  $M_{500} \sim 2 \times 10^{12} M_{\odot}$ , vanishing at  $10^{11} M_{\odot}$  and present but smaller above  $10^{13} M_{\odot}$ . At lower masses, the effect is likely due to the decreasing significance of AGN feedback relative to star formation feedback with decreasing mass, even while star formation feedback remains sufficient to cause a deviation from self-similarity. At higher masses, gravity has greater ability to counter the effect of feedback.

## 5. Future Work

We will extend this work to the simulation study of the CGM using other multiwavelength probes. In addition to the tSZ effect measurement, X-ray observation is another tool to characterize the hot gas properties around the galaxy groups and clusters (e.g., Anderson et al. 2015; van de Voort et al. 2016; Singh et al. 2018; Li 2020). Similar to the  $Y$ – $M$  relation, the soft X-ray luminosity of the galaxy clusters is expected to follow the mass-scaling relation. Several observations suggest that the emission is affected by the nongravitational heating, such as the AGN feedback (e.g., Pratt et al. 2009; Planck Collaboration et al. 2011; Anderson et al. 2015). Using the soft

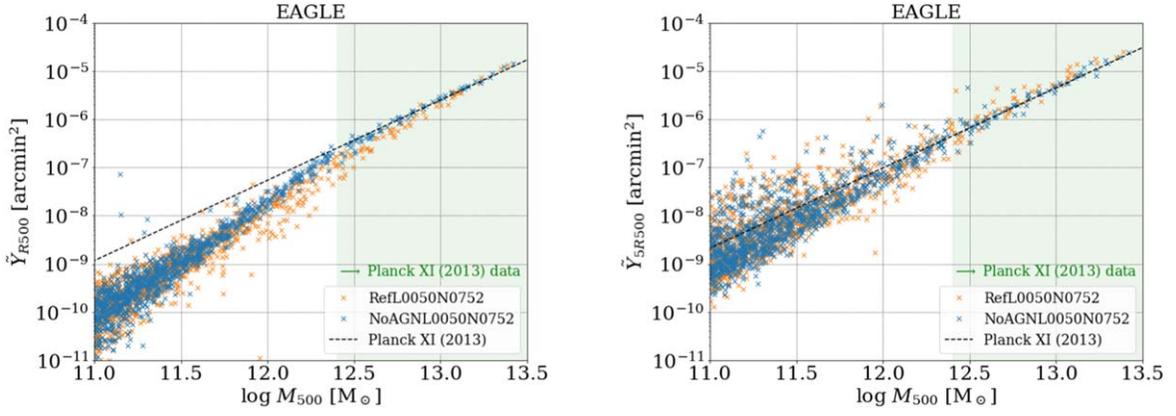
X-ray luminosity allows us to study the mass-dependent heating mechanism around the galaxies.

Besides the tSZ and X-ray imaging, several observational tools have been emerging to explore the hot halo across the electromagnetic spectrum. They include X-ray spectroscopy, which can probe the gas metallicity and thus the chemical composition of the CGM (Vijayan & Li 2022), and fast radio bursts (FRBs), which provide the line-of-sight free electron density with its frequency-dependent dispersion measure (e.g., Prochaska & Zheng 2019; Ravi 2019; Macquart et al. 2020; Connor & Ravi 2021). Traditional UV absorption-line spectroscopy can also trace the hot gas via high-energy ions such as O VI and Ne VIII. However, none of these methods can individually link the corresponding observation to a complete set of physical properties of the CGM, such as density, temperature, and metallicity. The best opportunity to gather the spatial and thermodynamic information of the hot CGM will be available through combining the observations across multiple wavelengths. This comprehensive approach will allow us to compare the relative importance of the feedback mechanisms that affect the regulation of star formation at various lengths and mass scales. In future work, we will address how the CGM can be used to distinguish different feedback models and how the relative importance of those models changes with the physical properties of galaxies. We will also study what combinations of the observational probes give the best understanding of the gas properties to break the degeneracy among them.

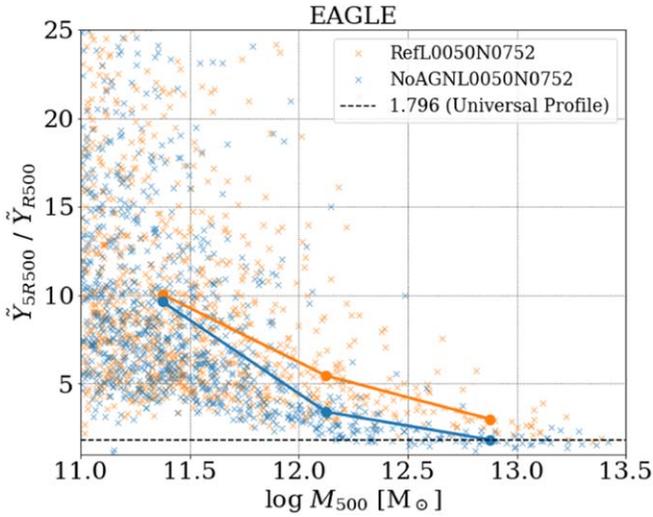
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*Software:* `astropy` (Astropy Collaboration et al. 2013), `numpy` (Harris et al. 2020), `matplotlib` (Hunter 2007), `Mop-c GT` (“Model-to-observable projection code for Galaxy Thermodynamics”),<sup>17</sup> `scipy` (Virtanen et al. 2020).

<sup>17</sup> <https://github.com/samodeo/Mop-c-GT>



**Figure 7.** The tSZ flux integrated out to  $R_{500}$  (left) and  $5R_{500}$  (right) for the galaxy samples in EAGLE Ref-L0050N0752 (blue crosses) and NoAGN-L0050N0752 (orange crosses) simulations. The PCXI best-fit scaling relation is shown as a dotted line. The shaded green region shows the mass range of the galaxy samples used in the PCXI analysis.



**Figure 8.**  $Y_{5R_{500}}/Y_{R_{500}}$  of the EAGLE galaxies in Ref-L0050N0752 and NoAGN-L0050N0752. The dashed line shows the conversion factor of 1.796 for the UPP. The blue (Ref-L0050N0752) and orange (NoAGN-L0050N0752) dots connected by the solid lines indicate the median value of the ratios in each  $M_{500}$  mass bin. The mass bins are separated by 0.75 dex between  $10^{11.0} M_{\odot}$  and  $10^{13.25} M_{\odot}$ .

## Appendix A

### Computing tSZ Signal Using Particle Data of Cosmological Simulations

Using particle data of the cosmological simulations, we calculate a total Comptonization parameter within a certain radius by computing each particle's tSZ signal strength. First, we calculate the electron number density

$$n_e = \frac{\rho_{\text{gas}}}{\mu_e m_{\text{H}}}. \quad (\text{A1})$$

The tSZ signal of a particular gas cell,  $\Upsilon$ , is

$$\begin{aligned} \Upsilon &= n_e \sigma_T \frac{k_B T_e}{m_e c^2} \\ &= \left( \frac{m_{\text{gas}}}{\mu_e m_{\text{H}}} \right) \sigma_T \frac{k_B T_e}{m_e c^2}, \end{aligned} \quad (\text{A2})$$

where  $\mu_e$  is the mean molecular weight per free electron in units of hydrogen mass (i.e., amu). Then, we integrate the

signals for all the gas particles within the radius, for example,  $R_{500}$ ,

$$Y_{R_{500}} D_A^2(z) = \sum \Upsilon(r < R_{500}). \quad (\text{A3})$$

The TNG data set gives the internal energy of each gas cell, and we need to convert it to a gas temperature. The mean molecular weight  $\mu$  there is defined as the average mass per particle:

$$\mu = \frac{(\sum_j n_j m_j) + n_e m_e}{(\sum_j n_j) + n_e} \frac{1}{m_p}, \quad (\text{A4})$$

where  $n_j$  and  $m_j$  are the number density of atoms and the corresponding mass, respectively. If we consider hydrogen ( $n_{\text{H}}$ ,  $m_p$ ), helium ( $n_{\text{He}}$ ,  $4m_p$ ), and all the other metals ( $n_z$ ,  $m_z$ ) with an approximation  $m_z \sim 4m_p$ , the total gas-mass density  $\rho_{\text{gas}} = n_{\text{H}} m_p + n_{\text{He}} (4m_p) + n_z m_z$  and

$$\begin{aligned} \mu &= \frac{n_{\text{H}} m_p + n_{\text{He}} (4m_p) + n_z m_z}{n_{\text{H}} + n_{\text{He}} + n_z + n_e} \frac{1}{m_p} \\ &= \frac{1}{\left[ n_{\text{H}} + n_{\text{He}} + \left( \frac{\rho_{\text{gas}}}{m_z} - \frac{1}{4} n_{\text{H}} - n_{\text{He}} \right) + n_e \right] / \rho_{\text{gas}}} \frac{1}{m_p} \\ &= \frac{4}{1 + 3X + 4Xx_e}, \end{aligned} \quad (\text{A5})$$

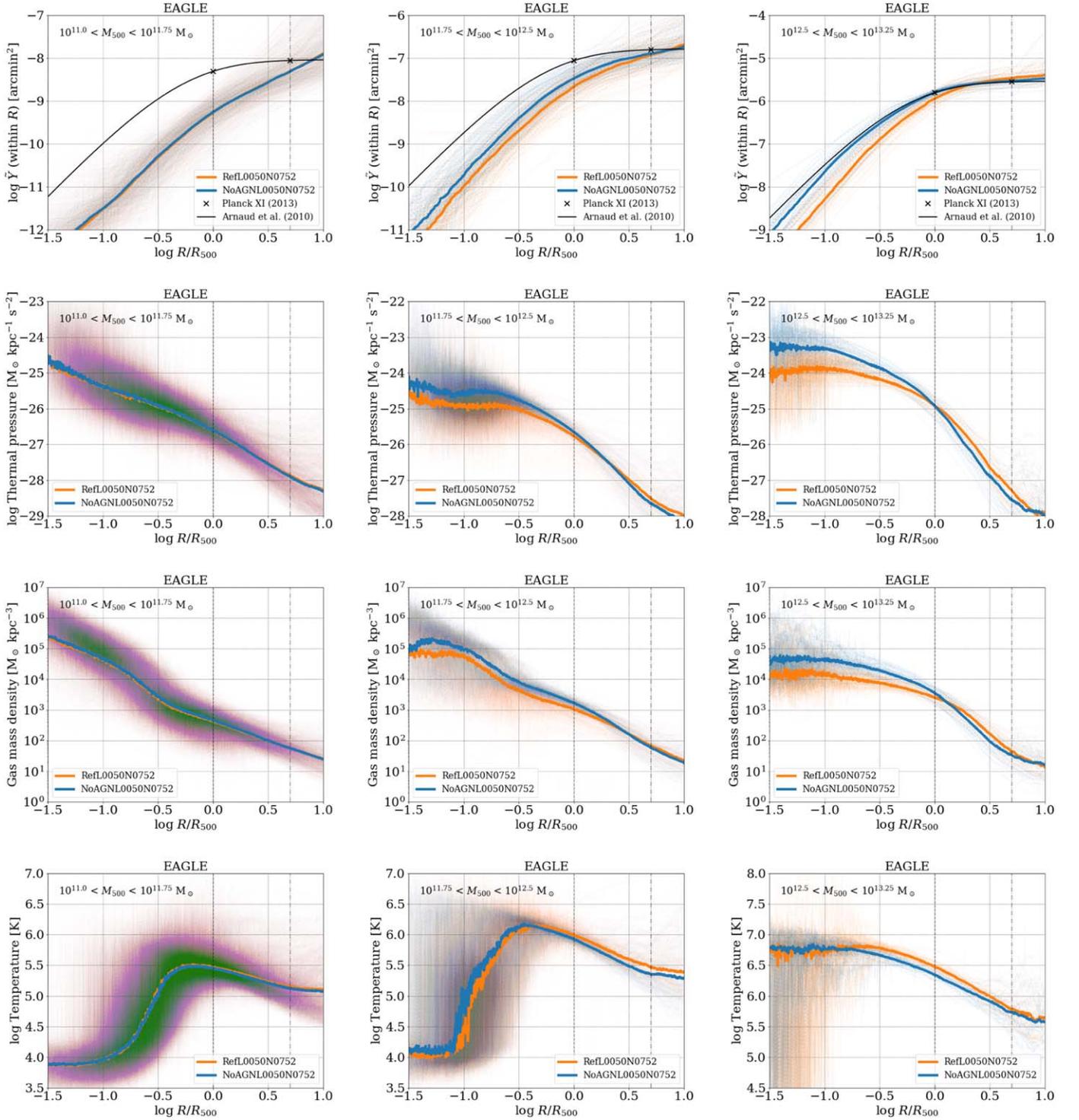
where  $x_e = n_e/n_{\text{H}}$  is the electron abundance.  $X = (n_{\text{H}} m_p)/\rho_{\text{gas}} = 0.76$ . (The EAGLE simulation does not provide electron abundance, so we assume fully ionized gas following the approach of Lim et al. 2018 for the simulation.) Then, the (specific) internal energy of the cell

$$u = \frac{1}{\gamma - 1} \left( \frac{1}{\mu} \right) k_B T_e, \quad (\text{A6})$$

where the adiabatic index  $\gamma = 5/3$ , and the gas temperature of the cell is

$$T_e = (\gamma - 1) \frac{u}{k_B} \mu. \quad (\text{A7})$$

In Section 3.4, where we compare the radial tSZ profiles of  $\sim 10^{12} M_{\odot}$  halos, we present the thermal pressure profiles as

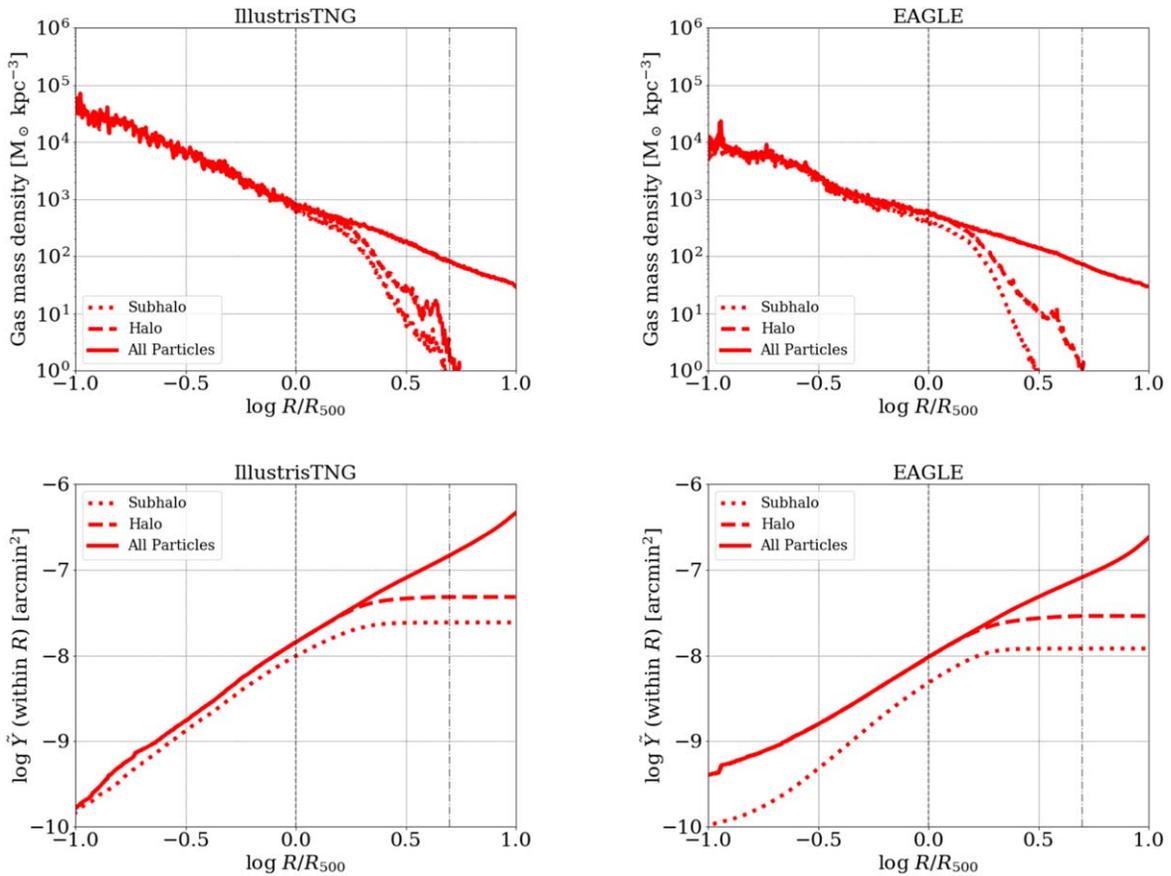


**Figure 9.** Radial profiles of  $10^{11.0} M_{\odot} < M_{500} < 10^{11.75} M_{\odot}$  (left column),  $10^{11.75} M_{\odot} < M_{500} < 10^{12.5} M_{\odot}$  (middle column), and  $10^{12.5} M_{\odot} < M_{500} < 10^{13.25} M_{\odot}$  (right column) galaxies in EAGLE Ref-L0050N0752 (orange) and NoAGN-L0050N0752 (blue) simulations. We show the cumulative tSZ flux, thermal pressure, gas-mass density, and temperature from top to bottom. The faint dotted lines are the radial profiles of the individual galaxies. The black crosses in the tSZ radial profile plots (top row) denote the tSZ signal at  $R_{500}$  and  $5R_{500}$  for the  $10^{11.375}$ ,  $10^{12.125}$ , and  $10^{12.875} M_{\odot}$  mass halos (left to right, respectively), inferred from the self-similar relation of PCXI. The mean  $R_{500}$  of each mass bin corresponds to 90, 160, and 300 physical kpc, respectively. The solid black lines in those plots show the spherically integrated tSZ flux as a function of radius, derived from the universal pressure profile (Arnaud et al. 2010), normalized by the PCXI measurement.

well. Note that the electron pressure  $P_e$  and the thermal pressure  $P_{\text{th}}$  are related by

$$P_e \simeq \left( \frac{2 + 2X}{3 + 5X} \right) P_{\text{th}} \quad (\text{A8})$$

(Battaglia et al. 2017), and we can directly calculate the thermal pressure from the particle data using the gas mass, internal energy per unit mass, and the volume of the gas cell.



**Figure B1.** Average radial gas-mass density (upper) and tSZ flux (lower) of the  $\sim 10^{12} M_{\odot}$  galaxies in TNG and EAGLE. Dotted, dashed, and solid lines show the radial profiles calculated using particles within a subhalo, a halo group, and the entire particles within the simulation volume around the center of the galaxy, respectively. The figures indicate that the discrepancy in the radial profiles exists depending on the particle selection in both simulations. It implies that the inclusion of the two-halo term is essential in comparing the simulation models to the observations, particularly at large radii from the center. Note that we show the profiles to contrast the particle selections in deriving the radial profiles. Readers need to avoid the physical interpretation comparing the radial profiles of the simulations.

## Appendix B

### Effect of Neighboring Particles in Calculating Radial Profiles

In the reconstruction of radial profiles and  $Y$  versus  $R$  in this paper from TNG and EAGLE, we do not limit ourselves to the particles in the halo or subhalo being considered but rather include all particles, that is, neighboring halos, subhalos, and the particles not bound to any halo/subhalo are included. This appendix justifies the choice. The large-volume TNG and EAGLE simulation data sets provide two kinds of group catalogs. Those are the halo (group) and subhalo (galaxy) catalogs derived by the FoF (Huchra & Geller 1982; Davis et al. 1985) and subfind algorithms (Springel et al. 2001), respectively. Simulation studies often limit the analysis to the particles associated with a single halo or subhalo. However, as shown in Figure 5, the effect of the “two-halo” term becomes not negligible at large radii from the center of the galaxy (Vikram et al. 2017; Hill et al. 2018; Moser et al. 2021).

Figure B1 shows the average radial gas-mass density and tSZ flux profiles in the TNG (TNG100-3) and EAGLE (Ref-L0050N0752) simulations. We randomly sampled  $\sim 50$  Milky Way-sized, quiescent galaxies in each simulation and calculated the tSZ flux and gas-mass density profiles similar to Figure 6. The dotted, dashed, and solid lines in the figure show the average radial profiles calculated using particle data affiliated with a single (central) subhalo, the parent FoF halo,

and the entire particle data in the simulation volume, respectively. As illustrated in Figure 2 of Moser et al. (2021), where they showed the gas-mass density and thermal pressure profiles, the two-halo term starts to contribute to the radial profile beyond  $R_{200} \sim 1.5R_{500}$ . It is clear that the  $Y$  versus  $R$  and radial profiles would show a striking and nonintuitive slope change beyond  $R_{500}$  where particles from neighboring halos are not included. It would be impossible to reproduce the PCXI results without including this effect, which is sensible because on-sky SZ measurements cannot distinguish between one halo and its neighbors. Even at  $R_{500}$ , one begins to see a distinction in the radial profiles, and to a lesser extent in  $Y$  versus  $R$ , considering only the central subhalo rather than the larger halo in which it resides. The figures make it clear that consideration of the CGM outskirts requires the inclusion of neighboring halos (i.e., all particles).

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