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GLADE+ : an extended galaxy catalogue for multimessenger searches with advanced gravitational-wave detectors

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ABSTRACT

We present GLADE+, an extended version of the GLADE galaxy catalogue introduced in our previous paper for multimessenger searches with advanced gravitational-wave detectors. GLADE+ combines data from six separate but not independent astronomical catalogues: the GWGC, 2MPZ, 2MASS XSC, HyperLEDA, and WISExSCOSPZ galaxy catalogues, and the SDSS-DR16Q quasar catalogue. To allow corrections of CMB-frame redshifts for peculiar motions, we calculated peculiar velocities along with their standard deviations of all galaxies having *B*-band magnitude data within redshift $z = 0.05$ using the ‘Bayesian Origin Reconstruction from Galaxies’ formalism. GLADE+ is complete up to luminosity distance $d_L = 47_{-2}^{+4}$ Mpc in terms of the total expected *B*-band luminosity of galaxies, and contains all of the brightest galaxies giving 90 per cent of the total *B*-band and *K*-band luminosity up to $d_L \simeq 130$ Mpc. We include estimations of stellar masses and individual binary neutron star merger rates for galaxies with *W1* magnitudes. These parameters can help in ranking galaxies in a given gravitational wave localization volume in terms of their likelihood of being hosts, thereby possibly reducing the number of pointings and total integration time needed to find the electromagnetic counterpart.

Key words: catalogues – galaxies: distances and redshifts.

1 INTRODUCTION

During their third observing run (O3), the network of advanced gravitational-wave (GW) detectors consisting of the two interferometers of Advanced LIGO (Aasi et al. 2015), Advanced Virgo (Acernese et al. 2014), and the recently joined KAGRA (Akutsu et al. 2021) issued tens of prompt public alerts of significant GW events to allow for follow-up observations with electromagnetic (EM) and neutrino observatories (Abbott et al. 2019b; LIGO Scientific Collaboration & Virgo Collaboration 2020). Each alert includes the posterior probability distribution of the source sky position, for which the 90 per cent credible localization area is typically a few hundred square degrees large (Abbott et al. 2018a). For compact binary coalescence (CBC) events, luminosity distances can also be inferred, and thus full 3D sky localization maps are released (Abbott

et al. 2019b). Such events include suspected binary neutron star (BNS) mergers, which are the leading candidates for joint GW + EM observations (see e.g. Abbott et al. 2017b).

An effective and typical way to find an EM counterpart for a GW event is to target a ranked set of potential host galaxies within the 3D localization volume with follow-up telescope observations, and to look for variations between time-separated images of them (Abadie et al. 2012; Nissanke, Kasliwal & Georgieva 2013; Fan, Messenger & Heng 2014; Hanna, Mandel & Vousden 2014; Gehrels et al. 2016; Singer et al. 2016). Besides telescope-dependent technical considerations and the probability distribution of the source location, one can also take into account astrophysically motivated factors in the ranking of host candidates. For example when dealing with a BNS merger event, one can sort host galaxy candidates by a model-based estimation of the BNS merger rate in them, which is expected to correlate with the star formation rate (Phinney 1991) (and thus with the blue luminosity, see e.g. Baret et al. 2012) and/or with the stellar mass (Artale et al. 2019, 2020) of the galaxies. In the favourable case

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of identifying the host galaxy of a GW event through observing its EM counterpart (see e.g. Abbott et al. 2017c; Coulter et al. 2017), one can refine the parameter estimations for the GW source with priors derived from the EM counterpart and properties of the host (see e.g. Gao et al. 2017; Guidorzi et al. 2017; Abbott et al. 2019a), or use GW parameter estimations to draw conclusions on the EM source (e.g. Abbott et al. 2018b; Rezzolla, Most & Weih 2018; Radice & Dai 2019), the host (e.g. Adhikari et al. 2020), or on cosmological parameters including the Hubble constant (see Abbott et al. 2017a, 2021a, b). Host galaxy candidates associated with detected CBC events can also be utilized in dark siren measurements of the Hubble constant when no EM counterpart of one or more CBC GW events is found (Fishbach et al. 2019; Soares-Santos 2019; Gray et al. 2020), in mapping out the expansion history (Mukherjee & Wandelt 2018; Bera et al. 2020; Diaz & Mukherjee 2021; Mukherjee et al. 2021a), and in testing the general theory of relativity based on GW propagation (Mukherjee, Wandelt & Silk 2020, 2021b).

In Dálya et al. (2018), we introduced the *Galaxy List for the Advanced Detector Era* (GLADE) value-added full-sky catalogue of galaxies (Dálya et al. 2016), which since then has extensively been used by the LIGO-Virgo Collaboration (Abbott et al. 2017a, 2020b, 2021b,c) and others (see Dálya et al. 2018 for references, and e.g. Fishbach et al. 2019; Lundquist et al. 2019; Antier et al. 2020; Gompertz et al. 2020; Finke et al. 2021; Keivani et al. 2021; Paterson et al. 2021; Sasada et al. 2021) for the purposes mentioned above. Additionally, GLADE has been integrated into tools supporting optimal selections of target galaxies for follow-up observations (Coughlin et al. 2019; Rana & Mooley 2019; Salmon et al. 2020; Wyatt et al. 2020; Xu et al. 2020), and was used for identifying or extracting information on potential hosts of fast radio bursts (see e.g. CHIME/FRB Collaboration 2019), short gamma-ray bursts (see e.g. Dichiara et al. 2020), supernova shock breakouts (Alp & Larsson 2020), and other transient phenomena observed in X-ray (Caglar & Hudaverdi 2017), optical (see e.g. Andrew, Swihart & Strader 2021; Killestein et al. 2021), and radio bands (see e.g. Riseley et al. 2017; Kuiack et al. 2021).

In this paper, we introduce an updated, extended, and improved version of GLADE, which we will refer to as GLADE+. GLADE+ contains ~ 22.5 million galaxies and $\sim 750\,000$ quasars (compared to ~ 3 million galaxies and $\sim 300\,000$ quasars in GLADE) as a result of cross-matching GLADE with the WISE \times SuperCOSMOS Photometric Redshift Catalogue (WISE \times SCOSPZ) and updating the quasar data base of GLADE with the SDSS-DR16Q quasar catalogue. To allow corrections of CMB-frame redshifts for peculiar motions, we calculated peculiar velocities along with their standard deviations of all galaxies having B -band magnitude data within redshift $z = 0.05$ using the ‘Bayesian Origin Reconstruction from Galaxies’ formalism (Mukherjee et al. 2021c). We estimated stellar masses and BNS merger rates for all GLADE+ galaxies with $W1$ -band magnitude data, and included these in the catalogue. The aim of these improvements is to continue serving the purposes mentioned above and potentially serve new ones, especially in light of the ongoing preparations for the fourth observing run of the LIGO-Virgo-KAGRA network scheduled for 2022 (Abbott et al. 2018a).

This paper is organized as follows. In Section 2, we introduce the catalogues and methods we used to create GLADE+, including the peculiar velocity correction and the stellar mass and BNS merger rate estimations. In Section 3, we quantify the completeness of GLADE+ based on the measured B - and K_s -band luminosities of galaxies. In Section 4, we describe the exact format of the catalogue, and in Section 5 we draw our conclusions.

Throughout this paper, we adopt a flat Λ cold dark matter cosmology with the following parameters from the Planck 2018 results: $H_0 = 100h = 67.66 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3111$, and $\Omega_\Lambda = 0.6889$ (Planck Collaboration VI 2020).

2 CONSTRUCTION OF THE GLADE+ CATALOGUE

We have constructed GLADE+ from six separate but not independent¹ astronomical catalogues: the Gravitational Wave Galaxy Catalogue² (GWGC, see White, Daw & Dhillon 2011), HyperLEDA³ (Makarov et al. 2014), the Two Micron All-Sky Survey Extended Source Catalog⁴ (2MASS XSC, see Jarrett et al. 2000 and Skrutskie et al. 2006), the 2MASS Photometric Redshift Catalog⁵ (2MPZ, see Bilicki et al. 2014), the WISE \times SCOS Photometric Redshift Catalogue⁶ (WISE \times SCOSPZ, see Bilicki et al. 2016a), and the Sloan Digital Sky Survey quasar catalogue from the 16th data release⁷ (SDSS-DR16Q, see Lyke et al. 2020). As we have used the first four of these catalogues in creating the GLADE galaxy catalogue, the relevant characteristics of these are summarized in Dálya et al. (2018). Note that there are several other survey catalogues available offering deep digital observations, such as the DESI Legacy Survey (Dey et al. 2019), Pan-STARRS (Flewelling 2020, Beck et al. 2021), or SkyMapper (Onken et al. 2019), which we plan to incorporate in future versions of the catalogue. In Section 2.1, we only describe the WISE \times SCOSPZ and SDSS-DR16Q catalogues, and discuss the cross-matching between GLADE and WISE \times SCOSPZ, as well as the results we obtained. In Section 2.2, we describe the method we used for estimating peculiar velocities, and in Section 2.3 we introduce the methods used to estimate the stellar masses of and BNS merger rates in the individual galaxies.

2.1 Cross-matching

The WISE \times SCOSPZ catalogue was constructed by cross-matching the AllWISE full-sky release (Cutri et al. 2014) of the Wide-field Infrared Survey Explorer (WISE, see Wright et al. 2010), which is the most comprehensive survey of the mid-infrared sky, and the SuperCOSMOS Sky Survey (Hambly et al. 2001), the result of an automated scanning and digitizing of photographic plates from the United Kingdom Schmidt Telescope and the Palomar Observatory Sky Survey-II. WISE \times SCOSPZ contains ~ 20 million galaxies with photometric redshifts calculated using an artificial neural network algorithm (Collister & Lahav 2004). The redshifts have errors nearly independent of distance, with an overall accuracy of $\sigma_z/(1+z) \simeq 0.033$ (Bilicki et al. 2016a). The catalogue contains magnitude information in the B_J and R_F bands from SuperCOSMOS (Peacock et al. 2016) and in the $W1$ and $W2$ WISE bands. In order to calculate the magnitudes of the galaxies in the Johnson–Cousins B band, we used the colour equations presented in Peacock et al. (2016).

¹2MPZ was created by cross-matching 2MASS XSC, WISE, and SuperCOSMOS and using an artificial neural network approach trained on several redshift surveys to derive the photometric redshifts, hence it is not independent from these catalogues; for details see Bilicki et al. (2014).

²<http://vizier.u-strasbg.fr/viz-bin/VizieR?-source=GWGC>

³<http://leda.univ-lyon1.fr/>

⁴<https://old.ipac.caltech.edu/2mass/>

⁵<http://ssa.roe.ac.uk/TWOMPZ.html>

⁶<http://ssa.roe.ac.uk/WISExSCOS.html>

⁷https://www.sdss.org/dr16/algorithms/qso_catalog/

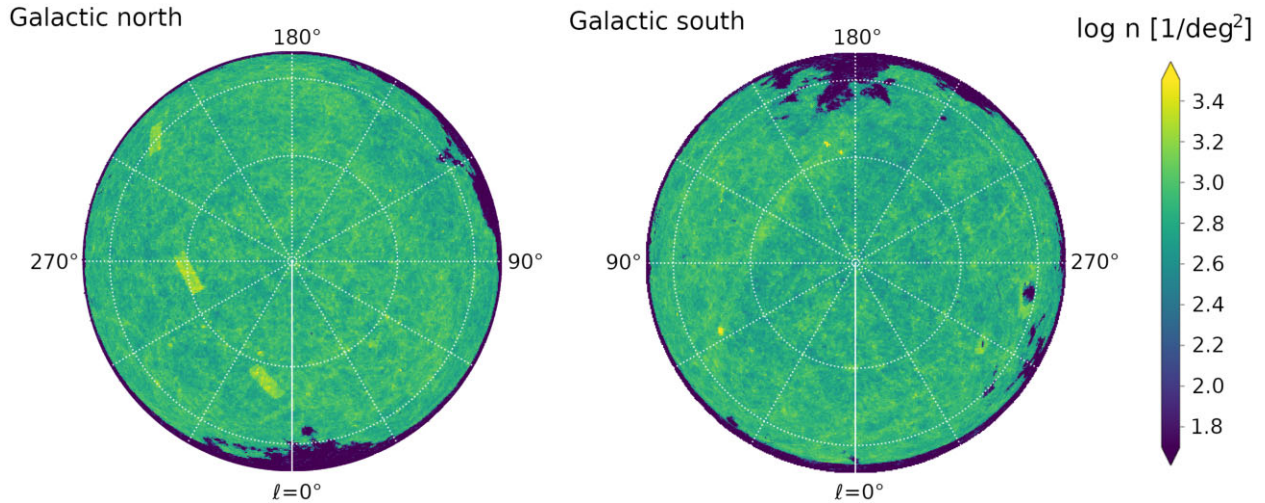


Figure 1. The base-10 logarithm of the number density (n) of objects in GLADE+ using azimuthal projection with galactic coordinates. The plane of the Milky Way obscures the visibility of background galaxies, hence the (blue) underdense regions at the edges of the plots. The overdense (yellow) patches and stripes originating from the HyperLEDA catalogue show up as a result of deeper, more sensitive surveys (such as SDSS and GAMA) that have been made towards the corresponding sky locations.

We have created GLADE+ by cross-matching GLADE v2.4⁸ with the WISExSCOSPZ catalogue and then replacing the quasars with the newer set from SDSS-DR16Q and removing the globular clusters. We could not use the method described in Dály et al. (2018) for cross-matching GLADE with WISExSCOSPZ, as duplicate galaxies could not simply be found by their designations. Hence, we used a resolution of 2 arcsec, i.e. if a WISExSCOSPZ galaxy lied closer to a GLADE galaxy than this threshold, we treated them as being the same object and merged them. This distance threshold was motivated by the fact that 2 arcsec is the maximal resolution in the WISExSCOSPZ catalogue and false positive associations start to dominate above this value. The order of magnitude of the threshold is also consistent with that of previous GLADE cross-matches; note however that in previous applications we could take other parameters of the galaxies into account as well, such as luminosity distances and B magnitudes (Dály et al. 2018).

The GLADE catalogue incorporated the SDSS-DR12Q quasar catalogue, which, in GLADE+, we replaced entirely with the more recent and extended SDSS-DR16Q catalogue. This catalogue contains data for $\sim 750\,000$ quasars (including the $\sim 300\,000$ quasars published in SDSS-DR12Q), which makes it the largest selection of spectroscopically confirmed quasars to date.

Cross-matching and updating the catalogues resulted in the GLADE+ catalogue containing 23 181 758 objects from which 22 431 348 are galaxies and 750 410 are quasars. The sky distribution of GLADE+ objects are shown in Fig. 1 as a density plot. The plane of the Milky Way is clearly noticeable in the figure, as the gas and dust reduces the visibility towards those directions and different sky surveys used various cuts in galactic latitude. Other anisotropies are arising from the different sensitivities and footprints of the various sky surveys.

Due to incorporating the WISExSCOSPZ catalogue, 21 165 400 galaxies in GLADE+ had $W1$ magnitudes available from that sample, which we used to estimate their stellar masses and the BNS merger rates in them, see Section 2.3. In order to supplement as many of the

remaining ones with $W1$ as possible, we first cross-matched them with AllWISE, using 3 arcsec ching radius and keeping only the closest matches. We found a match for over 1.76 million galaxies, which left us with $\sim 260\,000$ without a WISE counterpart. For part of the latter (32 153 galaxies) we provide $W1$ apparent magnitude estimates based on K_s -band measurements available in 2MASS. To obtain these estimates, we first fitted the mean $W1 - K_s$ colour as a function of redshift for sources with available spectroscopic redshifts, K_s and $W1$ magnitudes in the 2MPZ catalogue. This effective $W1 - K_s$ colour relation is then added to the K_s -band measurements for those sources in GLADE+ which have K_s magnitudes from 2MASS but do not have the $W1$ ones, to obtain the $W1$ magnitude estimates.

2.2 Peculiar velocity correction

The correction of peculiar velocities for low-redshift galaxies is essential to correctly calculate the true redshift. The estimation of peculiar velocities for galaxies in GLADE+ is made using the method proposed by Mukherjee et al. (2021c), which relies on a Bayesian formalism called ‘Bayesian Origin Reconstruction from Galaxies’ (BORG). The BORG forward modelling method infers a probabilistic and physically plausible model of the three-dimensional cosmic matter distribution from observed galaxies in cosmological surveys to derive the linear and partially the non-linear component of the velocity field (see e.g. Jasche & Wandelt 2013; Jasche, Leclercq & Wandelt 2015; Lavaux & Jasche 2016; Jasche & Lavaux 2019). This method solves a large-scale Bayesian inverse problem by fitting a dynamical structure formation model to data, and estimates the initial conditions of the early Universe from which presently observed structures can be explained. The BORG algorithm marginalizes over unknown galaxy bias and accounts for selection and evolutionary effects while providing the velocity field as part of the dynamical model. The reliability of this method is verified with an N -body simulation to check the consistency of the velocity fields provided by the posterior distributions sampled by BORG and the one provided by the original N -body simulation (Mukherjee et al. 2021c). This method gives a set of points in the parameter space (a spatial grid of 256^3 values with a spatial resolution of $2.64 \text{ Mpc } h^{-1}$ for the

⁸The specifics of different GLADE versions are described on the GLADE website: <http://glade.elte.hu>

initial conditions plus the bias parameters) that provides a numerical approximation of the posterior distribution of these parameters given the observed large-scale structure observation. For each sample of the posterior, the initial and final positions of the dark matter particles are provided, from which the velocity field can be estimated using the Simplex-in-Cell estimator (SIC, Hahn, Angulo & Abel 2015; Leclercq et al. 2017). More details on this method and its validation can be found in Mukherjee et al. (2021c).

Along with the velocity estimation from BORG, we also include the contribution from the non-linear virial component of the velocity field. The radial components of the virial velocities are modelled as Gaussian random variates with variance (Sheth & Diaferio 2001; Mukherjee et al. 2021c)

$$\sigma_{\text{vir}} = 476 g_v (\Delta_{\text{nl}}(z)E(z)^2)^{1/6} \left(\frac{M_{\text{h}}}{10^{15} M_{\odot} h^{-1}} \right)^{1/3}, \quad (1)$$

where $g_v = 0.9$, $\Delta_{\text{nl}}(z) = 18\pi^2 + 60x - 32x^2$, and $x = \Omega_{\text{m}}(1+z)^3/E^2(z) - 1$; $E(z)$ is the cosmological expansion function. In order to use this relation, we need to estimate the halo mass M_{h} . We have used a mass–luminosity relation (Vale & Ostriker 2004)

$$L_B = \frac{A(M_{\text{h}}/M_r)^b}{[c + (M_{\text{h}}/M_r)^{dk}]^{1/k}}, \quad (2)$$

where $A = 5.7 \times 10^9$, $M_r = 10^{11} M_{\odot}$ is the parameter capturing the knee in the mass–luminosity relation, $b = 4$, $c = 0.57$, $d = 3.72$, and $k = 0.23$, to estimate the mass from the B -band luminosity L_B of the galaxy. The total variance in the velocity field is then calculated as $\sigma_{\text{tot}}^2 = \sigma_{\text{BORG}}^2 + \sigma_{\text{vir}}^2$.

Using this method we have estimated the mean value⁹ of the velocity field to all GLADE+ galaxies which cross-match with the 2M+ + compilation (Lavaux & Hudson 2011; Mukherjee et al. 2021c) for galaxies up to redshift $z = 0.05$ for which the B -band luminosity is available, along with the standard deviation σ_{tot} of the peculiar velocity which includes both linear and non-linear components of the velocity field. The conversion from the heliocentric redshift to the CMB-frame redshift are performed using the observation of the CMB temperature anisotropy by FIRAS (Fixsen et al. 1996). The uncertainty in the velocity error is translated into an uncertainty in the redshift and is provided in the catalogue.

2.3 Stellar mass and BNS merger rate estimations

According to Artale et al. (2019), stellar masses (i.e. the total mass of the active and remnant stars) of galaxies strongly correlate with the merger rates of the colliding binaries. These parameters can help in ranking the galaxies in a GW localization volume for EM follow-up observations, thereby possibly reducing the number of pointings and the total integration time needed to find the EM counterpart. The stellar mass is also one of the key parameters of the formation and evolution of galaxies (e.g. Engler et al. 2020; Ahad et al. 2021; van Loon, Mitchell & Schaye 2021). Hence, we aimed to estimate stellar masses and CBC merger rates for as many galaxies in GLADE+ as possible.

To estimate the stellar mass of a galaxy, the so-called (stellar) mass-to-light ratio (M_*/L) is required that can be obtained from stellar population synthesis models. Then the estimation can be performed with spectral energy distribution fitting, or based on one or more magnitude bands or colours (see Courteau et al. 2014 for a

review). Since only a few magnitude bands (B and some infrared) are contained in GLADE+, and no spectral energy distributions are available from the source catalogues, it is straightforward to use one of the magnitude-based stellar mass estimation methods. Note that spectral energy distributions could be obtained from external catalogues, with which more precise stellar mass estimations could be given; this is something we consider to do in later versions of our catalogue. The mid-infrared M_*/L is relatively insensitive to the different stellar populations, particularly in the absence of ongoing star formation, and in addition it is not very sensitive to dust attenuation (see e.g. Wen et al. 2013; Röck et al. 2015). Therefore, we estimated the stellar masses based on the mid-infrared WISE magnitudes.

According to Kettley et al. (2018) the mass-to-light ratio in the $W1$ band is

$$M_*/L_{W1} = 0.65 \pm 0.07 \quad (3)$$

for passive galaxies, which, according to the authors' claim, can provide at least as accurate mass estimates for galaxies with redshifts $z \leq 0.15$ as other more complex methods. For the galaxies with active star formation, they give the next equation:

$$\log_{10}(M_*/L_{W1}) = -0.4 \pm 0.2. \quad (4)$$

Here, the stellar mass is in the usual solar mass (M_{\odot}) unit, and the WISE magnitudes are in the Vega system. We separated the passive and active galaxies: $W2 - W3 \leq 1.5$ indicates active star formation (Jarrett et al. 2013; Cluver et al. 2014). We identified 799 703 passive and 18 351 034 active galaxies. There are 2628 585 galaxies having no galaxy type; we assumed that they are active. We calculated the stellar masses using equations (3) and (4) for galaxies having WISE $W1$ magnitude according to their galaxy type. The $W1$ luminosities can be calculated as

$$L_{W1}(L_{\odot}) = 10^{-0.4(M - M_{\odot, W1})}, \quad (5)$$

where $M_{\odot, W1} = 3.24^{\text{m}}$ is the $W1$ magnitude of the Sun and

$$M = m + 5 - 5 \log_{10} d_L - K \quad (6)$$

is the absolute magnitude of the galaxy in the $W1$ band, m is the apparent $W1$ magnitude, d_L is the luminosity distance and

$$K = -7.1 \log_{10}(1+z) \quad (7)$$

is the K -correction of Kettley et al. (2018), where z is the heliocentric redshift. We do not apply correction for extinction because WISE magnitudes are already corrected for that (Bilicki et al. 2016b). We only accepted stellar masses larger than $10^5 M_{\odot}$ as this value is the lower limit of the stellar masses of dwarf (e.g. Garrison-Kimmel et al. 2019). (There were no galaxies having stellar mass smaller than $10^5 M_{\odot}$.) Using this method, we could estimate the stellar mass of 21 779 322 (~ 97 per cent) of the galaxies in GLADE+.

We also provide the error of the stellar mass of each galaxy using the propagation of uncertainty. The errors of luminosity distances were calculated from the errors of redshifts in the cases where it is not known. Where the error of the $W1$ magnitude was not known for an individual galaxy, we used the mean $W1$ error from WISExSCOSPZ instead. As a result of our calculations, we have found that the mean (median) relative stellar mass error is 40 (33) per cent for passive galaxies and 67 (58) per cent for active galaxies. The codes of the stellar mass estimation are publicly available on the GLADE website.

⁹The direction of the velocity field is chosen such that the positive value of the velocity field indicates that the object is moving away from us.

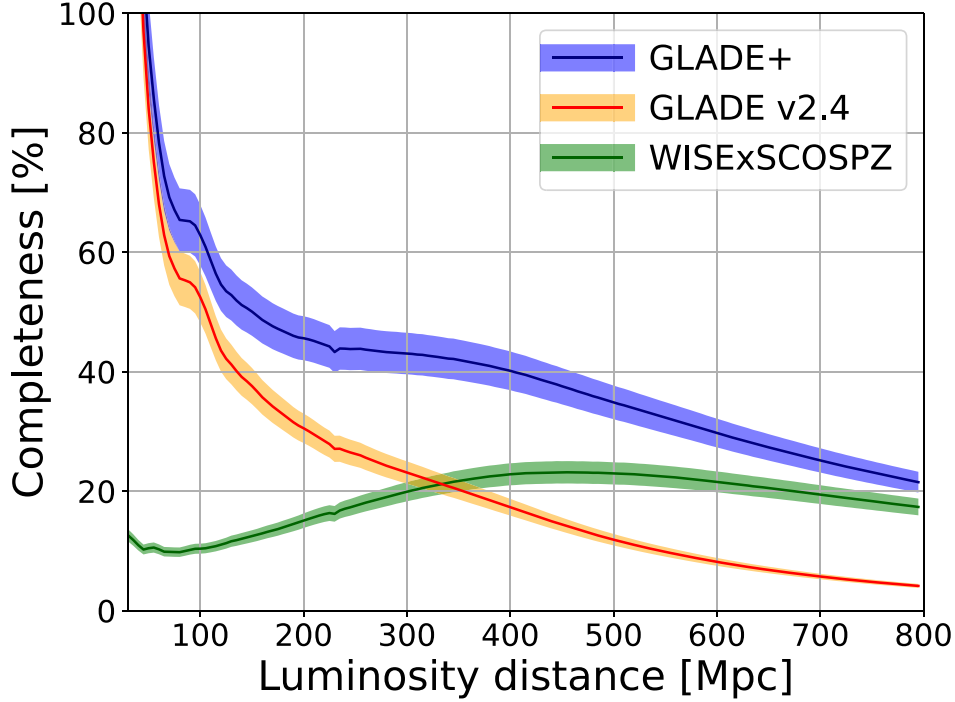


Figure 2. The completeness in terms of the normalized integrated B -band luminosity of galaxies in GLADE+ (blue) and its constituent catalogues (i.e. GLADE v2.4 shown in orange and WISExSCOSPZ shown in green) within luminosity distances indicated on the x -axis. The normalization is carried out with the integrated B -band luminosity calculated from an average B -band luminosity density of a complete catalogue of homogeneously distributed galaxies (see Kopparapu et al. 2008 for details). The completeness value exceeds 100 per cent within 47 Mpc due to a local overdensity of galaxies around the Milky Way. The completeness below ~ 330 Mpc is mostly due to GLADE galaxies, and WISExSCOSPZ galaxies contribute more for distances larger than ~ 330 Mpc.

The BNS merger rates (N_{BNS}) can be calculated from the stellar mass values according to Artale et al. (2019):

$$\log_{10}(N_{\text{BNS}}/\text{Gyr}) = (1.15 \pm 0.08) \log_{10}(M_*/M_{\odot}) - (7.22 \pm 0.22), \quad (8)$$

which is valid below $z \leq 0.1$ and for galaxies with stellar masses $M_* > 10^7 M_{\odot}$, so we calculated the BNS merger rates for galaxies satisfying these conditions. Applying these criteria resulted in 3156 544 galaxies with BNS merger rates. In addition, Artale et al. (2019) provide equations for the merger rates of binary black holes:

$$\log_{10}(N_{\text{BBH}}/\text{Gyr}) = (0.80 \pm 0.07) \log_{10}(M_*/M_{\odot}) - (4.14 \pm 0.19), \quad (9)$$

and of black hole–neutron star pairs as well:

$$\log_{10}(N_{\text{NSBH}}/\text{Gyr}) = (0.87 \pm 0.08) \log_{10}(M_*/M_{\odot}) - (4.99 \pm 0.22). \quad (10)$$

Note that we only provide merger rates for BNSs in GLADE+, however using the stellar masses from the catalogue and equations (9)–(10), the binary black hole and the black hole–neutron star merger rates can be calculated as well.

3 CATALOGUE COMPLETENESS

Following the methods we introduced in Dályá et al. (2018), we quantify the completeness of GLADE+ using two different methods: (i) by comparing, within different d_L limits, the integrated B luminosity of GLADE+ galaxies to calculated reference values, and

(ii) by comparing luminosity distributions of GLADE+ galaxies to the Schechter function within different luminosity distance shells.

The first method was originally used by White et al. (2011) to calculate the completeness of the GWGC catalogue. Here, we compare the integrated B -band luminosity of GLADE+ galaxies within different luminosity distance limits to the total B -band luminosity we would expect from the same volume given a complete catalogue of homogeneously distributed galaxies with B -band luminosity density $(1.98 \pm 0.16) \times 10^{-2} L_{10} \text{Mpc}^{-3}$, where $L_{10} = 10^{10} L_{B,\odot}$ and $L_{B,\odot}$ is the solar luminosity in the B band. Fig. 2 shows a comparison between the completeness values inferred using this method for GLADE+ and its constituent catalogues, GLADE v2.4 and WISExSCOSPZ. Completeness values over 100 per cent are results of local overdensities of galaxies. The drop in completeness around $d_L \simeq 220$ Mpc corresponds to the distance limit for our peculiar velocity correction (see Section 2.2). As we can see from Fig. 2, most of the completeness below ~ 330 Mpc comes from GLADE v2.4, and at larger distances contributions from WISExSCOSPZ galaxies start to dominate. Based on this completeness measure, GLADE+ is complete up to $d_L = 47^{+4}_{-2}$ Mpc. GLADE+ has a completeness of ~ 55 per cent within the single-detector LIGO Livingston BNS range during O3 (130 Mpc) and ~ 45 per cent within the maximal planned single-detector BNS range in O4 (190 Mpc, see Abbott et al. 2020a). Note that the BNS range is the average distance from which a GW detector can detect a circular binary of two $1.4 M_{\odot}$ neutron stars with a signal-to-noise ratio of 8, where the average is calculated over all possible sky positions and orbital inclinations, the maximum distance from which a BNS (i.e. with optimal sky direction and inclination) can be detected is ~ 2.26 times larger (Finn & Chernoff 1993). Furthermore, the data-driven projections obtained by Petrov

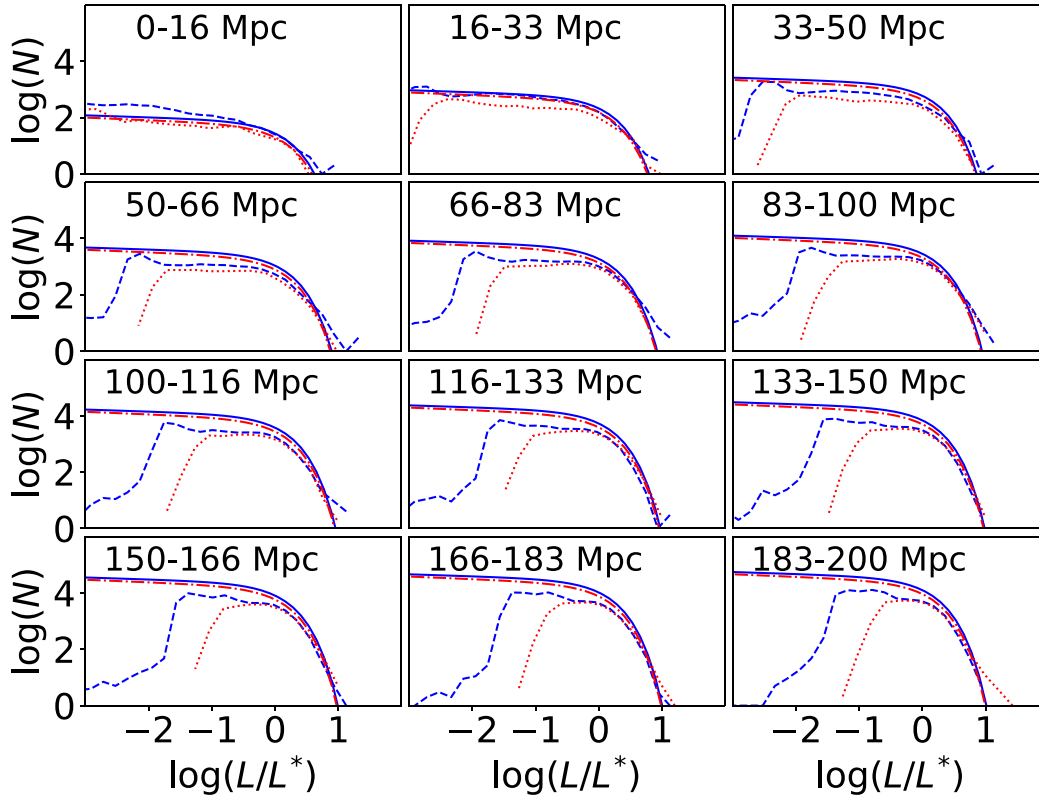


Figure 3. Luminosity histograms of GLADE+ galaxies within different luminosity distance shells in terms of their measured B -band and K_s -band luminosities (blue dashed and red dotted lines, respectively), compared to the same histograms we expect for complete catalogues based on B - and K_s -band Schechter functions (blue solid and red dash-dotted lines, respectively). L^* is the characteristic luminosity of the Schechter function.

et al. (2022) suggest that the median luminosity distance of BNS mergers observed during O4 will be 352.8 Mpc.

We have also used a second method to characterize the completeness of GLADE+. Here, we divided the galaxies into 12 luminosity distance shells, each having a width of $\Delta d_L = 16.7$ Mpc. We have constructed histograms of B - and K_s -band luminosities of GLADE+ galaxies for each shell, which we show in Fig. 3 together with their corresponding Schechter functions. For the B -band Schechter function we have used the following values from Gehrels et al. (2016): $\phi^* = 1.6 \times 10^{-2} h^3 \text{Mpc}^{-3}$, $M_B^* = -20.47$, and $a = -1.07$. For the K_s -band Schechter function we have used the following parameters from Kochanek et al. (2001): $\phi^* = 1.16 \times 10^{-2} h^3 \text{Mpc}^{-3}$, $M_K^* = -23.39$, and $a = -1.09$. Fig. 3 shows that as distance increases, more and more faint galaxies are missing from GLADE+ in both bands. We can also see that the faint limit of our catalogue decreases more rapidly in the K_s band. Note that as Fig. 3 shows, even though the Schechter function fits quite well for the galaxy distributions in both the B and the K_s bands, it seems to systematically underestimate the number of the brightest galaxies. Hence, the completeness measure we can derive from it can only be an approximation of the catalogue’s true completeness.

We have compared the integrated luminosity of a subset of galaxies giving 90 per cent of the total luminosity in each shell to the expected value corresponding to the Schechter function. In this analysis, we have increased the luminosity distance limit to $d_L = 500$ Mpc and the shell width to $\Delta d_L = 20$ Mpc. The completeness of GLADE+ in the B and K_s bands in the different bins is shown in Fig. 4. The figure shows that the completeness for this subset of galaxies decreases more rapidly in the K_s band over $d_L \simeq$

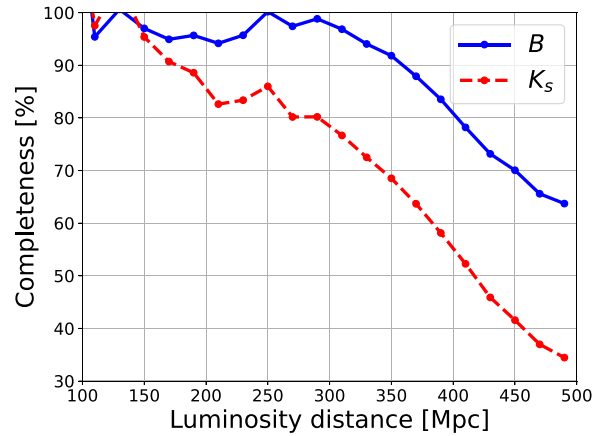


Figure 4. Completeness of GLADE+ in various distance shells having a 20 Mpc width relative to the B -band (blue solid line) and the K_s -band Schechter functions (red dashed line) for the galaxies giving 90 per cent of the total luminosity in the given band. Note that even if the completeness is 100 per cent or above at a given distance, GLADE+ can still lack fainter galaxies.

100 Mpc, and using this definition GLADE+ is complete up to ~ 130 Mpc in the B and K_s bands. Note that we have considered only 90 per cent of all the galaxies for this analysis, so even if the completeness is 100 per cent or above at a given distance, GLADE+ can still lack a large number of the faintest galaxies there.

4 DESCRIPTION OF THE GLADE+ CATALOGUE

The GLADE+ catalogue is available as a `txt` file on the GLADE website.¹⁰ Columns of the file contain the following data (where available) for each GLADE+ object:

- (1) GLADE+ catalogue number
- (2) Principal Galaxies Catalogue number
- (3) Name in the GWGC catalogue
- (4) Name in the HyperLEDA catalogue
- (5) Name in the 2MASS XSC catalogue
- (6) Name in the WISExSuperCOSMOS catalogue (wiseX)
- (7) Name in the SDSS-DR16Q catalogue
- (8) Object type flag: ‘*Q*’ means that the source is from the SDSS-DR16Q catalogue, ‘*G*’ means that it is from another catalogue and has not been identified as a quasar
- (9) Right ascension in degrees (J2000)
- (10) Declination in degrees (J2000)
- (11) Apparent *B* magnitude
- (12) Absolute error of apparent *B* magnitude
- (13) *B* magnitude flag: ‘0’ if the *B* magnitude is measured, ‘1’ if it is calculated from the *B_J* magnitude
- (14) Absolute *B* magnitude
- (15) Apparent *J* magnitude
- (16) Absolute error of apparent *J* magnitude
- (17) Apparent *H* magnitude
- (18) Absolute error of apparent *H* magnitude
- (19) Apparent *K_s* magnitude
- (20) Absolute error of apparent *K_s* magnitude
- (21) Apparent *W1* magnitude
- (22) Absolute error of apparent *W1* magnitude
- (23) Apparent *W2* magnitude
- (24) Absolute error of apparent *W2* magnitude
- (25) *W1* flag: ‘0’ if the *W1* magnitude is measured, ‘1’ if it is calculated from the *K_s* magnitude
- (26) Apparent *B_J* magnitude
- (27) Absolute error of apparent *B_J* magnitude
- (28) Redshift in the heliocentric frame
- (29) Redshift converted to the Cosmic Microwave Background (CMB) frame
- (30) Redshift correction flag: ‘0’ if the CMB frame redshift and luminosity distance values given in columns 29 and 33 are not corrected for the peculiar velocity, ‘1’ if they are corrected values
- (31) Error of redshift from the peculiar velocity estimation
- (32) Measurement error of heliocentric redshift
- (33) Luminosity distance in Mpc units
- (34) Error of luminosity distance in Mpc units
- (35) Redshift and luminosity distance measurement flag: ‘0’ if the galaxy has no measured redshift or distance value, ‘1’ if it has a measured photometric redshift from which we have calculated its luminosity distance, ‘2’ if it has a measured luminosity distance value from which we have calculated its redshift, ‘3’ if it has a measured spectroscopic redshift from which we have calculated its luminosity distance
- (36) Stellar mass in $10^{10} M_{\odot}$ units
- (37) Absolute error of stellar mass in $10^{10} M_{\odot}$ units
- (38) Stellar mass flag: ‘0’ if the stellar mass was calculated assuming no active star formation, ‘1’ if the stellar mass was calculated assuming active star formation

- (39) Base-10 logarithm of estimated BNS merger rate in the galaxy in Gyr^{-1} units
- (40) Absolute error of estimated BNS merger rate in the galaxy

5 CONCLUSIONS

The GLADE+ galaxy catalogue is an extended version of the GLADE catalogue we have optimized for multimessenger searches with advanced GW detectors. It contains more than 23 million objects from which more than 22 million are galaxies and $\sim 750\,000$ are quasars. As the effects of peculiar motions are important for nearby galaxies for both the EM follow-up and cosmological analyses, we have estimated the peculiar velocities along with their standard deviations using the BORG forward modelling method for galaxies in GLADE+ with *B*-band magnitude data and having redshifts $z \leq 0.05$. GLADE+ is complete up to $d_L = 47_{-2}^{+4}$ Mpc in terms of the cumulative *B*-band luminosity of galaxies, and contains all of the brightest galaxies giving half of the total *B*-band (*K_s*-band) luminosity up to $d_L \simeq 250$ Mpc ($d_L \simeq 390$ Mpc).

As according to theoretical models the stellar masses of galaxies strongly correlate with the merger rates of colliding binaries, we have calculated the stellar masses and the BNS merger rates (together with their errors) of each galaxy having WISE magnitudes. These parameters can help in ranking the galaxies in a given GW localization volume for EM follow-up observations, thereby possibly reducing the number of pointings and the total integration time needed to find the EM counterpart.

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¹⁰GLADE website: <http://glade.elte.hu>

¹¹<https://www.aquila-consortium.org/>

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DATA AVAILABILITY

The data underlying this article are available at the GLADE website, <http://glade.elte.hu/>.

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