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The delay time distribution of Type-Ia supernovae in galaxy clusters: the impact of extended star-formation histories

Jonathan Freundlich^{1,2★} and Dan Maoz^{1★}

¹*School of Physics and Astronomy, Tel Aviv University, Tel Aviv 69978, Israel*

²*Observatoire Astronomique, Université de Strasbourg, CNRS, 11 rue de l'Université, F-67000 Strasbourg, France*

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ABSTRACT

The delay time distribution (DTD) of Type-Ia supernovae (SNe Ia) is important for understanding chemical evolution, SN Ia progenitors, and SN Ia physics. Past estimates of the DTD in galaxy clusters have been deduced from SN Ia rates measured in cluster samples observed at various redshifts, corresponding to different time intervals after a presumed initial brief burst of star formation. A recent analysis of a cluster sample at $z = 1.13\text{--}1.75$ confirmed indications from previous studies of lower redshift clusters, that the DTD has a power-law form, $\text{DTD}(t) = R_1(t/\text{Gyr})^\alpha$, with amplitude R_1 , at delay $t = 1$ Gyr, several times higher than measured in field-galaxy environments. This implied that SNe Ia are somehow produced in larger numbers by the stellar populations in clusters. This conclusion, however, could have been affected by the implicit assumption that the stars were formed in a single brief starburst at high z . Here, we re-derive the DTD from the cluster SN Ia data, but relax the single-burst assumption. Instead, we allow for a range of star-formation histories and dust extinctions for each cluster. Via MCMC modelling, we simultaneously fit, using stellar population synthesis models and DTD models, the integrated galaxy-light photometry in several bands, and the SN Ia numbers discovered in each cluster. With these more-realistic assumptions, we find a best-fitting DTD with power-law index $\alpha = -1.09_{-0.12}^{+0.15}$, and amplitude $R_1 = 0.41_{-0.10}^{+0.12} \times 10^{-12} \text{ yr}^{-1} M_\odot^{-1}$. We confirm a cluster-environment DTD with a larger amplitude than the field-galaxy DTD, by a factor $\sim 2\text{--}3$ (at 3.8σ). Cluster and field DTDs have consistent slopes of $\alpha \approx -1.1$.

Key words: methods: statistical – stars: evolution – transients: supernovae – galaxies: clusters: general – galaxies: star formation.

1 INTRODUCTION

Type-Ia supernovae (SNe Ia) are explosions involving the thermonuclear combustion of white-dwarf stars, although the precise identity of the exploding systems, the processes leading to explosion, and the development of the explosion itself, are all far from clear yet (see Maoz, Mannucci & Nelemans 2014; Livio & Mazzali 2018, for reviews). Together with their more-numerous cousins, core-collapse supernovae (CC-SNe), SNe Ia are the production factories for the elements from oxygen to the iron-peak elements, they are also sources of kinetic-energy feedback that plays a role in galaxy formation, and their remnants are the acceleration sites of cosmic rays. Among the observable physical properties of the SN Ia population, the delay-time distribution (DTD) of SNe Ia has been useful for investigating both the SN Ia progenitor question and the contribution of SNe Ia to the chemical and structural evolution of galaxies. As its name implies, the DTD is the distribution of times elapsed between the formation, at time zero, of a hypothetical stellar population, of unit mass, and the explosion, as SNe Ia, of some of its white-dwarf remnants (see Maoz & Mannucci 2012, for a review

on the DTD and its measurement). The DTD, usually expressed as an SN Ia rate per formed stellar mass as a function of delay time, embodies the ‘impulse response’ of the SN Ia phenomenon, and can provide clues relevant for studying all of the issues mentioned above.

The DTD is directly linked to the lifetimes of SN Ia progenitor systems, and it could, in principle, vary with environment, cosmic time, metallicity, and initial mass function (IMF). The past decade has seen numerous efforts to measure the DTD, using diverse observational methodologies and analyses, applied to SNe Ia in different environments and times (field galaxies, cluster galaxies, low and high redshift). For field galaxies, the DTD has been estimated by deconvolving the volumetric SN Ia rate, as a function of redshift, from the cosmic star formation history (SFH), where the DTD is the convolution kernel. Similarly, the field-galaxy DTD has been also derived from comparison of SN Ia rates in galaxies to their individual galaxy SFHs. Many of these studies have converged toward a DTD with a power-law time dependence, t^α , with $\alpha \sim -1$, between delays $t \sim 40$ Myr and a Hubble time. The DTD is assumed to be zero at $0 \leq t \lesssim 40$ Myr, since a stellar population has not yet formed any white dwarfs, essential for an SN Ia explosion, during that interval. Observations have yet to constrain well the DTD at short delays, of the order of ~ 100 Myr and below. The latest analyses of SN Ia rates in field galaxies (e.g. Maoz & Graur 2017) find a DTD with a

* E-mail: jonathan.freundlich@mail.huji.ac.il (JF); maoz@astro.tau.ac.il (DM)

power-law index $\alpha = -1.1 \pm 0.1$, and a time-integrated (from 40 Myr to 13.7 Gyr) number of SNe Ia per formed stellar mass,¹ $N_{\text{Ia}}/M_* = (1.3 \pm 0.1) \times 10^{-3} M_{\odot}^{-1}$.

In galaxy clusters, whose stars are assumed to have been formed in single brief bursts at redshifts $z_f \sim 3 - 4$ (more on this assumption further below), the DTD has been recovered simply by measuring the SN Ia rate in galaxy-cluster samples as a function of their redshift, with different redshifts corresponding to different delays since the epoch of the single star-formation burst. Most recently, Friedmann & Maoz (2018, FM18) searched for SNe Ia in a sample of clusters at $z = 1.13-1.75$ monitored by the *Hubble Space Telescope* (HST). Fitting a power-law DTD to the SN Ia rate in this high- z sample jointly with the rates measured by previous studies of lower redshift samples, FM18 found best-fitting parameters $\alpha = -1.30_{-0.16}^{+0.23}$ (i.e. a power-law index somewhat, although not significantly, steeper than seen in field environments) and a Hubble-time-integrated SN Ia production efficiency, N_{Ia}/M_* , at least several times larger than the value of N_{Ia}/M_* measured by field-SN Ia studies. Both of these results confirmed indications deduced from previous, lower- z , cluster samples (e.g. Maoz, Sharon & Gal-Yam 2010; Barbary et al. 2012).

The high normalization of the DTD of SNe Ia in galaxy clusters is particularly intriguing. It implies that something in the mode of star-formation or stellar evolution that took place in these environments is conducive to an overall larger production of SNe Ia, at all delays. FM18 speculated that, at the root of the phenomenon, could be, e.g. the formation of a higher fraction of close binaries in the mass range of stars that are the progenitors of SNe Ia; or perhaps a ‘middle-heavy’ IMF, with an excess of intermediate-mass stars, that evolve into white dwarfs in the mass range that explode as SNe Ia. There is already mounting evidence for IMF peculiarities in massive elliptical galaxies, the very same type that dominate galaxy clusters, even if there is yet disagreement about the driving agents of these peculiarities (galaxy mass, metallicity, age, etc.). van Dokkum & Conroy (2010), Treu et al. (2010), Conroy & van Dokkum (2012), Conroy, van Dokkum & Villaume (2017), Cappellari et al. (2012), Cappellari et al. (2013), Lyubenova et al. (2016), Davis & McDermid (2017) and others have deduced the existence of a ‘bottom-heavy’ IMF in such galaxies: the IMF slope, rather than becoming shallower at stellar masses below $\sim 0.5 M_{\odot}$, continues with the ‘Salpeter’ slope down to the hydrogen-burning limit at $\sim 0.1 M_{\odot}$. The stellar atmospheres in such galaxies also have a high ratio of α elements to iron $[\alpha/\text{Fe}]$, compared to the stars in lower mass and disc galaxies, and compared to the disc stars in the Milky Way (e.g. Conroy, Graves & van Dokkum 2014). Since the bulk of the α elements are produced by CC-SNe from massive stars, while iron is synthesized both in SNe Ia and CC-SNe (with roughly equal contributions to the universal iron budget, see Maoz & Graur 2017), a peculiar $[\alpha/\text{Fe}]$ ratio is therefore suggestive of IMF variations at the high end of the IMF: a change in the high-mass IMF slope might change the resulting mix of different CC-SN types, thus changing the integrated $[\alpha/\text{Fe}]$ from CC-SNe; and/or a change in the ratio of high-mass stars (that explode as CC-SNe) and intermediate-mass stars (some of whose white-dwarf descendants produce SNe Ia), could also affect $[\alpha/\text{Fe}]$. Hallakoun &

Maoz (2020) have recently discovered a Salpeter-like bottom-heavy IMF in yet another ancient, high $[\alpha/\text{Fe}]$, environment – the ‘metal-poor’ or ‘blue’ halo of the Milky Way, thought to be the stellar debris of a galaxy accreted by our Galaxy 10 Gyr ago. Although the low-mass end of the IMF is unlikely to be relevant for SNe Ia, whose white-dwarf progenitors derive from intermediate-mass main-sequence stars, these latest results show that IMF variations are possible among different galaxy populations or environments. The high-normalized DTD measured in galaxy clusters could thus be another indicator of an IMF that is peculiar to star formation in early-type galaxies, cluster environments, or early cosmic epochs. It could also be related to the long-known excess of the iron-to-stellar mass ratio in clusters, compared to expectations from SN yields (e.g. see Ghizzardi et al. 2020, for a recent analysis).

However, the measurement of the DTD in galaxy clusters, and the results described above, have hinged on the assumption that the stars in galaxy clusters were all formed in a short burst at high redshift, with negligible ensuing star forming thereafter. This assumption has permitted to associate directly the SN Ia rate per unit mass, measured in a sample of clusters at redshift z , with the DTD at a delay t_{delay} corresponding to the time elapsed between the formation redshift, z_f , and the cluster redshift z . The short-burst assumption is likely justified, at least when analysing clusters at $z \lesssim 1$. Spectral analyses of the stellar populations of cluster galaxies have long shown that the bulk of their stars were formed quickly at $z_f \sim 3-4$ (e.g. Daddi, Cimatti & Renzini 2000; Stanford et al. 2005; Eisenhardt et al. 2008; Snyder et al. 2012; Stalder et al. 2013; Andreon et al. 2014; Andreon, Dong & Raichoor 2016). Most recently, Salvador-Rusiñol et al. (2019) have analysed the ultraviolet-to-optical absorption lines in the high-signal-to-noise stacked spectra of tens of thousands of massive early-type galaxies at $z \sim 0.4$, and have set a stringent limit of < 0.5 per cent on the fraction of these galaxies’ stellar mass formed in the prior 2 Gyr.

A concern, however, for cluster galaxies observed at $z \sim 1-2$, corresponding to intervals of just $\sim 1-4$ Gyr after initial star formation, is that for such short delays the single-short-burst assumption might be unjustified, and could therefore result in an incorrectly determined DTD. Significant ongoing star formation in the cluster galaxies at the epoch of observation could, on the one hand, cause the DTD delay time effectively being probed by the measured SN Ia rate to be actually smaller than the delay implicit in the single short burst assumption, biasing high the DTD (since a shorter delay, in a monotonically decreasing DTD, would actually be probed). On the other hand, neglecting the presence of a young population, with a low mass-to-light ratio, in the cluster galaxies would lead to an overestimate of the formed stellar mass and to an underestimate of the SN Ia rate, and hence would bias the DTD low. To understand which, if either, of these biases are at work, and to obtain a more robust estimate of the DTD form and amplitude in clusters, the single-burst assumption must be relaxed.

In this paper, we re-analyse the SN Ia data in FM18 and the cluster SN Ia rates measured by previous studies, to find the observational constraints on the DTD in the presence of extended (rather than single-burst) star formation histories, but only those histories that are consistent with the observed luminosities and spectral energy distributions of the monitored clusters. To this end, we re-measure and revise the multiband near-infrared (rest-frame optical) photometry of the galaxy clusters in FM18. When modelling the data, we compare the predicted numbers of SNe Ia in every cluster in FM18 to the actual number of SNe Ia discovered in that cluster (rather than to the number of SNe Ia in the sample as a whole, as in FM18), thus utilizing more information. The SN Ia numbers are

¹As in most recent works on the subject, for the calculation of formed stellar mass, based on observed light from a stellar population, we will assume in this paper a Kroupa (2001) IMF. We note that this assumption is, more than anything, a calibration convention that facilitates comparisons of diverse measurements of SN Ia rates, that are all actually normalized relative to the observed stellar luminosity of their host galaxy populations, rather than to their current or formed mass, which is not directly observed.

also slightly revised, in view of recent follow-up observations of the FM18 sample that have led to some clarifications of SN classification and cluster membership of host galaxies. In addition we revise some of the estimates of formed stellar mass, and hence of the mass-normalized rates, in lower redshift cluster SN samples, based on improved modelling of their reported stellar luminosities.

2 REVISED CLUSTER STELLAR FLUXES, SUPERNOVA NUMBERS, AND SUPERNOVA RATES

2.1 Cluster stellar fluxes

FM18 measured and reported the net cluster–galaxy fluxes in three near-infrared broad photometric *HST* bands: F105W (centred at wavelength $\sim 1.05 \mu\text{m}$), F140W ($1.40 \mu\text{m}$), and F160W ($1.60 \mu\text{m}$). At the typical cluster redshifts, these bands correspond roughly to the rest-frame optical *B*, *V*, and *R* bands. In their analysis, FM18 used only the flux they measured in the reddest band flux, F160W, to estimate, under the early-short-burst assumption, the formed stellar mass of each cluster, by comparing the observed flux to predictions of spectral synthesis models of a single burst at a redshift $z_f = 3$ or 4. Without spectroscopy of all galaxies in the field of view, it is impossible to clearly discriminate cluster galaxies from foreground and background galaxies. Following previous studies, FM18 therefore performed a statistical subtraction of the foregrounds and the backgrounds. For this purpose, FM18 used *HST* images of the Hubble Ultra-Deep Field (HUDF; Beckwith et al. 2006), obtained with the same instrumental setup, in the role of the typical extragalactic sky in a non-galaxy-cluster direction. However, in the course of this work, we realized that the measurements of FM18 for some of the clusters in some of the bands are erroneous, apparently due to oversubtraction of foreground galaxy light. We have therefore remeasured the photometry of all of the FM18 clusters in their available bands.

For our new photometry, we first measured the total counts in a cluster image within a circular aperture, centred on a cluster’s brightest cluster galaxies (BCGs), with an aperture radius of 51 arcsec, but up to 55 arcsec for some of the clusters. The aperture radius for each cluster was chosen to avoid including within the aperture, if possible, bright foreground stars and galaxies near the aperture borders, whose subsequent subtraction would add noise to the photometric measurement. The aperture covers roughly 50 per cent of the WFC3-IR detector’s image area, comparable to the ‘fully time-covered area’ in each cluster field that was included in all *HST* epochs and all bands of a cluster, due to the image plane rotation from epoch to epoch in the course of the 2-yr SN Ia survey observations (see FM18). Only SNe Ia discovered within this area were included by FM18 in the SN Ia census counted towards the SN Ia rate, and therefore the light only from galaxies within this area needs to be integrated for estimating a cluster’s stellar luminosity and stellar mass that normalize the SN Ia rate. Next, we identified, within the aperture, foreground Galactic stars (easily identifiable in *HST* images via their diffraction spikes) and bright foreground galaxies that are brighter than a cluster’s BCGs. We photometered these stars and galaxies using small apertures (~ 1 arcsec, or as appropriate), subtracting the background counts estimated in annuli around the small apertures, and subtracted the foreground-star and bright-galaxy counts from the total counts of the large cluster aperture. Finally, we estimated the remaining non-cluster backgrounds and foregrounds by placing 13-arcsec radius apertures at diverse, relatively sparsely populated, locations around the image edges, avoiding stars and

obvious galaxy concentrations (either cluster- or foreground galaxies). These remaining backgrounds/foregrounds result from: residual dark current and residual cosmic ray events in the detector; airglow; zodiacal light; faint, both individually detected and undetected (below the flux detection limit) field (i.e. non-cluster) galaxies, and integrated extragalactic background light. The counts per pixel in these marginal-region apertures were used to estimate and subtract the total foreground and background counts within the large cluster aperture. The variations in counts from one edge region to another were used to estimate the uncertainty in the foreground/background subtraction, which dominates the uncertainty in the final cluster-light photometry. This method of estimating a cluster’s total stellar flux, by adding all counts within the cluster aperture, automatically includes all of the individually undetected faint cluster galaxies (i.e. galaxies with flux below the detection limit) plus any intracluster starlight, if it exists. Our new measurement procedure is essentially very similar to one of the two photometry approaches used by FM18, except for our field-edge-based foreground/background counts estimate, which replaces the HUDF-based estimate in FM18 and which, as noted, sometimes resulted in oversubtraction, and hence an underestimate of the cluster stellar fluxes. Table 1 lists our new photometric measurements, which supersede those published in table 5 of FM18.

2.2 Cluster SN Ia numbers

In their analysis of the *HST* SN survey data, FM18 discovered 29 transient events. Each transient’s multiband light curves were fit with SN Ia model curves, to gauge whether or not the transient was an SN Ia in the cluster. At the end of this process, FM18 reported that 11 events were likely SNe Ia in cluster galaxies, within the fully time-covered areas of the cluster fields, and a further four events were deemed possible cluster SNe Ia. At the time of this writing, the Supernova Cosmology Project (SCP) team, who proposed and led the *HST* observations carried out in 2014–2016, have not yet published the results of their own search for SNe in these data. However, Williams et al. (2020, W20) have reported the results of ground-based spectroscopic observations by the SCP of some of the host galaxies of the transients that the group discovered in the southern-hemisphere clusters in the sample. We have attempted to cross-match all of the transients mentioned in W20 with all of the transients previously reported by FM18. That attempt is summarized in Table 2. Some of the transients found by FM18 are not mentioned in W20 (even in those clusters for which W20 show images with the SCP-detected transients indicated), and vice versa. The non-mutual detections may have resulted from differing estimates, between the two groups, regarding the significance and reality of those detections. For example, some of the transients identified by W20 but not by FM18 are located near or on the nuclei of bright galaxies, some of the nuclei spectroscopically identified as active galactic nuclei (AGNs), which are often variable. None the less, the majority of the cluster SN Ia candidates are detected and similarly classified, based on their light curves, by the two studies. The spectroscopy of the host galaxies (and in one case, of the SN itself) by W20 can confirm or refute the cluster membership of some of the SN Ia candidates. We use the W20 results to slightly revise, for three clusters, the estimated FM18 numbers of SNe Ia that the clusters hosted. The revised SN Ia numbers are listed in Table 1, and supersede those in table 5 of FM18. The cases of clusters with re-assessed SN Ia numbers follow below.

In the cluster MOO1014, the galaxy spectroscopy by W20 leads them to revise the cluster redshift, from $z = 1.27$ down to $z = 1.23$. The event SNFM10, reported by FM18 and well-fit as an SN Ia in

Table 1. Revised cluster galaxy fluxes and hosted SN Ia numbers.

Cluster name (1)	z (2)	f_{F105W} (3)	f_{F140W} (4)	f_{F160W} (5)	N_{Ia} (6)	$N_{Ia?}$ (7)	τ_{vis} (8)	η_{-40} (9)	η_0 (10)	η_{+20} (11)
IDCS1426	1.75	0.50 ± 0.45	0.68 ± 0.10	0.53 ± 0.11	1	1	52, 89	0.71, 0.78	0.86, 0.87	0.75, 0.79
ISCS1432	1.40	0.70 ± 0.18	0.97 ± 0.06	0.81 ± 0.11	1	0	104, 87	0.88, 0.88	0.93, 0.96	0.86, 0.86
MOO1014	1.23 [†]	1.56 ± 1.20	3.02 ± 0.19	1.38 ± 0.17	2	0	98, 88	0.93, 0.92	0.99, 0.98	0.89, 0.89
MOO1142	1.19	1.94 ± 0.70	1.87 ± 0.13	–	0	0	89	0.87	0.93	0.88
SPARCS0224	1.63	1.00 ± 0.30	0.89 ± 0.16	0.67 ± 0.26	1	1	67	0.72	0.83	0.81
SPARCS0330	1.63	0.65 ± 0.65	0.76 ± 0.38	0.55 ± 0.15	0	1	112	0.77	0.87	0.85
SPARCS1049	1.70	1.30 ± 0.25	–	0.80 ± 0.14	0	0	41	0.56	0.63	0.62
SPARCS0035	1.34	1.28 ± 0.25	1.22 ± 0.18	–	0	1	75, 45	0.89, 0.88	0.95, 0.97	0.87, 0.88
SPT0205	1.58	0.90 ± 0.25	1.14 ± 0.13	1.09 ± 0.11	4	1	310	0.93	0.99	0.88
SPT2040	1.48	1.43 ± 0.35	2.13 ± 0.23	1.11 ± 0.30	3	0	101, 42	0.85, 0.85	0.96, 0.94	0.91, 0.84
SPT2106	1.13	3.55 ± 0.30	1.98 ± 0.13	–	1	0	90, 71	0.93, 0.95	0.98, 0.98	0.90, 0.91
XMM44	1.58	1.30 ± 0.10	1.32 ± 0.08	1.15 ± 0.13	1	0	70	0.69	0.79	0.77

Notes. (1) Cluster ID; (2) Cluster redshift; (3)–(5) Total measured galaxy cluster flux within the full-time-coverage area of the field in the F105W, F140W, and F160W bands, when available, in 10^{-16} erg cm $^{-2}$ s $^{-1}$ Å $^{-1}$; (6) and (7) Number of likely and possible cluster SNe Ia, respectively; (8) Cluster rest-frame monitoring time for each season, in days; (9)–(11) Detection-completeness factors for the periods before, during, and after each observing season, respectively, from FM18. When more than one season to a cluster, the two season’s monitoring times and completeness factors are listed. [†]Revised from $z = 1.27$, based on spectroscopy by Williams et al. (2020).

the cluster, is not mentioned in W20. Transient SNFM11, deemed by FM18 a cluster SN Ia, corresponds in W20 to SCP15C02. W20 find that its host is a foreground galaxy at $z = 0.7506$, and therefore we no longer count it as a cluster SN Ia. SNFM13 (=SCP14C01 in W20) was a hostless event marginally fit as an SN Ia, which FM18 classified as most probably a foreground event. However, W20 present a ‘live’ spectrum of the event itself and they conclude that it is likely an SN Ia at the cluster redshift. SCP15C04 is an event not reported by FM18, near the nucleus of its host galaxy, which may have affected the evaluation of its significance. W20 confirm the galaxy as a cluster member and list the event as a possible SN Ia. Given its non-detection by FM18 at its problematic location, and its uncertain classification by W20, we do not adopt this event into our cluster SN Ia sample. The remaining transients in cluster MOO104, whether reported by FM18, by W20, or by both, are all non-cluster on non-SNIa events (see Table 2). In the balance, one of this cluster’s assumed likely SNe Ia in the FM18 sample drops out (SNFM11), and one enters it as such (SNFM13). Together with SNFM10, the number of likely SNe Ia in this cluster remains unchanged at 2.

In SPARCS0224, FM18 discovered transient SNFM15 in a spectroscopically confirmed cluster host galaxy, which they classified as a possible SN Ia. W20 do not report this event. An additional transient, SNFM16 (=SCP16I02), was considered by FM18 a likely cluster SN Ia. W20 measure for its host a tentative redshift of $z = 1.62$, consistent with cluster membership, but they state that this redshift is not secure. For this cluster as well, the cluster SNe Ia numbers remain unchanged, at one possible and one likely SN Ia.

In cluster SPT0205, FM18 discovered two transients, SNFM22 (=SCP15A04) and SNFM23 (=SCP16A04) in the same host galaxy, separated by about a year in the observer frame. The light curves of both were well-fit by FM18 with SN Ia model light curves at the cluster redshift. A spectrum by W20 confirms the galaxy is in the cluster. Similarly, the host of SNFM24 (=SCP15A05), classified by FM18 as a cluster SN Ia, is confirmed by W20 as a cluster galaxy. For SNFM25 (=SCP15A01), also counted as a cluster SN Ia by FM18, the host galaxy was too faint for W20 to secure a redshift. In addition, and not detected by FM18, W20 report SCP15A06 as a likely SN Ia, with position nearly coincident with the nucleus of its host, which has a redshift $z = 1.3128$. This redshift is consistent with,

but somewhat lower than, the redshifts of the other host galaxies in this cluster measured by W20, $z = 1.322$ to $z = 1.340$, raising the possibility that this galaxy could be foreground to the cluster. All things considered, we count SCP15A06 as a ‘possible’ cluster SN Ia, in addition to the four likely SNe Ia already considered in FM18 in this cluster.

In the cluster SPT2040, FM18 discovered event SNFM26 (=SCP15E04 in W20) in a faint host galaxy. FM18 and W20 both consider this a likely cluster SN Ia event. No spectrum is reported by W20. Not detected by FM18, however, are two transients, SCP15E06 and SCP15E07, reported by W20 as likely SNe Ia, with their hosts confirmed as cluster members. We therefore add these two events to SNFM26, for a total of three likely SNe Ia in this cluster. We do not include in our tally an additional event in this cluster, SCP15E03, mentioned by W20 as a possible SN Ia, and coincident with the nucleus of a galaxy lacking a measured redshift.

For the cluster SPT2106, W20 report that SCP16D01 (not detected by FM18), deemed a likely SN Ia, is hosted by a galaxy at the cluster redshift. We therefore raise the number of likely SNe Ia in this cluster from zero to one. Finally, in cluster XMM44, FM18 classified SNFM29 (=SCP15G01 in W20) as a likely cluster SN Ia. The host spectrum by W20 confirms its cluster membership.

For the full FM18 cluster sample, the total number of cluster SNe Ia is thus revised from 11 to 14 likely SNe Ia, and from 4 to 5 possible SNe Ia.

2.3 Revised SN Ia rates

The observed SN Ia rate per formed stellar mass was calculated in FM18 as

$$R_{Ia,M} = \frac{N_{Ia,tot}}{\sum_i M_i \eta_i \tau_{vis,i}}, \quad (1)$$

where $N_{Ia,tot}$ is the total number of transients in a cluster sample that were classified as SNe Ia, M_i is the formed stellar mass under the assumption of a single instantaneous burst, $\tau_{vis,i}$ are the rest-frame monitoring times, and η_i are the detection-completeness factors. The sum in the denominator is simultaneously over the different clusters, their observing seasons, and the periods before, during, and after each season. As our revised numbers for the FM18 sample now

Table 2. Cross-match between transients in FM18 and W20.

Cluster	Name in FM18	Name in W20	Notes
IDCS1426	SNFM01	–	possible foreground SN Ia
	SNFM02	–	hostless, near field edge
	SNFM03	–	possible cluster SN Ia
	SNFM04	–	non-SN Ia
	SNFM05	–	foreground non-SN Ia
	SNFM06	–	likely cluster SN Ia
ISCS1432	SNFM07	–	likely cluster SN Ia
	SNFM08	–	hostless, near field edge
	SNFM09	–	not discovered in F140W search band
MOO1014	SNFM10	–	likely cluster SN Ia
	SNFM11	SCP15C02	foreground CC-SN at $z = 0.7506$
	SNFM12	SCP16C03	lensed background SN Ia at $z = 2.22$
	SNFM13	SCP14C01	likely cluster SN Ia, based on ‘live’ W20 spectrum
	–	SCP15C01	SN Ia in foreground at $z = 0.9718$
	–	SCP15C03	non-SN Ia, possibly $z = 1.01$
	–	SCP15C04	cluster host; W20: ‘possible SN Ia’
	–	SCP16C01	foreground non-SN Ia, $z = 9739$
–	SCP16C02	non-SN Ia, no host redshift	
MOO1142	SNFM14	–	hostless, non-SN Ia, peculiar light curve
SPARCS0224	SNFM15	SCP15I01	possible SN Ia in confirmed cluster member
	SNFM16	SCP15I02	likely SN Ia, host at cluster redshift but not secure
SPARCS0330	SNFM17	SCP16H01	possible cluster SN Ia
	SNFM18	–	non-SN Ia
SPARCS0035	SNFM19	–	hostless, non-SN Ia
	SNFM20	–	possible cluster SN Ia
SPT0205	SNFM21	SCP16A03	non-SN Ia
	SNFM22	SCP15A04	likely SN Ia in confirmed cluster galaxy
	SNFM23	SCP16A04	likely SN Ia in confirmed cluster galaxy
	SNFM24	SCP15A05	likely SN Ia in confirmed cluster galaxy
	SNFM25	SCP15A01	likely cluster SN Ia
	–	SCP15A02	non-SN Ia, $z = 0.8966$
	–	SCP15A03	cluster SN Ia, but near field edge
	–	SCP15A06	possible SN Ia, $z = 1.3128$, possibly cluster foreground
	–	SCP15A07	non-SN Ia at $z = 0.5015$
	–	SCP16A01	non-SN Ia, $z = 0.4998$
–	SCP16A02	non-SN Ia, $z = 0.7971$	
SPT2040	SNFM26	SCP15E04	likely cluster SN Ia
	SNFM27	SCP15E02	non-SN Ia
	–	SCP15E01	non-SN Ia, $z = 0.8398$
	–	SCP15E03	possible SN Ia, on nucleus of galaxy with no redshift
	–	SCP15E05	non-SN Ia, on nucleus of bright galaxy
	–	SCP15E06	likely SN Ia in confirmed cluster galaxy
	–	SCP15E07	likely SN Ia in confirmed cluster galaxy
	–	SCP15E08	possible SN Ia in broad-lined AGN at $z = 2.02$
–	SCP16E02	non-SN Ia, $z = 0.9435$	
SPT2106	SNFM28	–	non-SN Ia, near field edge
	–	SCP15D01	foreground SN Ia, $z = 0.5682$, near nucleus of bright galaxy
	–	SCP15D02	possible SN Ia on the nucleus of a cluster-member AGN
	–	SCP15D03	possible SN Ia near nucleus of possibly foreground galaxy at $z = 1.1130$
	–	SCP15D04	possible SN Ia, field edge, no clear host
	–	SCP16D01	likely SN Ia in confirmed cluster galaxy
	–	SCP16D02	non-SN Ia, near field edge
–	SCP16D03	non-SN Ia, $z = 0.6118$	
XMM44	SNFM29	SCP15G01	likely SN Ia in confirmed cluster galaxy

include 14 likely and 5 possible SN Ia, we adopt $N_{\text{Ia,tot}} = 16.5 \pm 2.5$ (systematic error) ± 4.1 (Poisson error).

To re-estimate the SNe Ia rate and error, as implied by the FM18 sample, we calculate $R_{\text{Ia,M}}$ with 10 000 realizations. In each realization, we draw a value of $N_{\text{Ia,tot}}$ from a distribution centred on its mean value, and having width according to its systematic and Poisson uncertainties.

In each realization and for each cluster, we also draw a delay time, uniformly spread within the range corresponding to formation redshifts between $z_f = 3 - 4$. The formed stellar mass is determined, as in FM18, i.e. by comparing the total stellar luminosity of a cluster in its reddest available band, and its uncertainty, to the luminosity in that band predicted by spectral-population synthesis (SPS) modelling, for a single burst given the delay time (the delay time range constituting a further source of uncertainty for the derived mass).

For the SPS generation we use the Pégase .3 SPS code (Fioc & Rocca-Volmerange 1997, 1999, 2019) instead of Starburst99 (Leitherer et al. 1999, 2010, 2014; Vázquez & Leitherer 2005) used by FM18. We assume a constant Solar stellar metallicity $Z = 0.020$, a Kroupa (2001) IMF, and Padova+AGB stellar evolution tracks (cf. Fioc & Rocca-Volmerange 2019, section 2.2.2 and references therein). We have verified that we obtain essentially identical model spectra with Pégase .3 and with Starburst99, for instantaneous bursts of star formation with the same input parameters.

For the full FM18 sample, we obtain an SNe Ia rate per unit mass of $R_{\text{Ia,M}} = 24.6 \pm 7.1 \times 10^{-14} \text{ yr}^{-1} M_{\odot}^{-1}$. We also rederive the SN Ia rate per unit rest-frame B -band stellar luminosity, $R_{\text{Ia,LB}}$, which is calculated the same as $R_{\text{Ia,M}}$, except that the formed stellar mass, M_i , is replaced by the B -band luminosity of each cluster, $L_{B,i}$, in equation (1). The B -band luminosity is deduced from the Pégase .3 spectrum. We find $R_{\text{Ia,LB}} = 0.64 \pm 0.19 \times 10^{-12} \text{ yr}^{-1} L_{\odot}^{-1}$, as listed in Table 3. The delay time range corresponding to a formation redshift $z_f = 3-4$ for the mean redshift of the sample is 2.2–2.8 Gyr.

Splitting the FM18 sample into two equal subsamples, one composed of the six lower- z clusters ($z = 1.1-1.4$) and the other of the six higher- z clusters ($z = 1.4-1.75$), the number of SN Ia in the lower- z sample is $N_{\text{Ia,low}} = 9.0 \pm 1.0$ (systematic) ± 3.0 (Poisson) and that in the higher- z sample $N_{\text{Ia,high}} = 7.5 \pm 1.5$ (systematic) ± 2.7 (Poisson). The corresponding SN Ia rates are respectively 19.5 ± 6.8 and $37.0 \pm 15.5 \times 10^{-14} \text{ yr}^{-1} M_{\odot}^{-1}$ per formed stellar mass, and 0.57 ± 0.20 and $0.74 \pm 0.31 \times 10^{-12} \text{ yr}^{-1} L_{\odot}^{-1}$ per unit B -band luminosity. The delay-time ranges probed by the two subsamples are respectively 2.7–3.2 and 1.8–2.4 Gyr. The revised full-sample measurements of the SN Ia rates and those for the two subsamples are all within $\lesssim 10$ per cent ($< 0.5\sigma$) of the FM18 values.

In the studies of cluster SN Ia rates previous to FM18, as compiled by Maoz & Graur (2017) and listed also in table 6 of FM18, cluster stellar masses were derived indirectly from the observed cluster luminosities. The colour of a cluster stellar population was either measured from multiband data, or assumed to be the colour of an old elliptical galaxy. Relations between mass-to-light ratio and colour, based on SPS calculations of an evolved single burst by Bruzual & Charlot (2003) were then used to convert the observed cluster luminosities to stellar masses. The stellar masses were in turn used to normalize the SN Ia rates per unit mass. To obtain more direct and consistent estimates of the rates per unit mass that are implied by previous cluster SN Ia studies, we have re-assessed the previous rates, finding the Pégase .3 model B -band mass-to-light ratio for each cluster sample, and using it to translate the rates per unit luminosity to rates per unit mass. As previously, the Pégase .3 models assume

Table 3. Revised SN Ia rates.

Source	z	t (Gyr)	$R_{\text{Ia,LB}}$ ($10^{-12} \text{ yr}^{-1} L_{B,\odot}^{-1}$)	$R_{\text{Ia,M}}$ ($10^{-14} \text{ yr}^{-1} M_{\odot}^{-1}$)
(1)	(2)	(3)	(4)	(5)
FM18 sample, single redshift bin				
FM18	1.35	2.5 ± 0.3	0.64 ± 0.19	24.6 ± 7.1
FM18 sample, two redshift bins				
FM18	1.58	2.1 ± 0.3	0.74 ± 0.31	37.0 ± 15.5
FM18	1.25	3.0 ± 0.3	0.57 ± 0.20	19.5 ± 6.8
Previous measurements				
B12	1.12	3.5 ± 0.3	$0.50^{+0.33}_{-0.28}$	$14.5^{+10.3}_{-8.8}$
GY02	0.90	4.4 ± 0.3	$0.80^{+1.06}_{-0.52}$	$18.1^{+24.1}_{-12.3}$
S10	0.60	5.9 ± 0.3	$0.35^{+0.31}_{-0.26}$	$5.8^{+5.5}_{-4.8}$
G08	0.46	6.9 ± 0.3	$0.31^{+0.51}_{-0.16}$	$4.5^{+7.4}_{-2.4}$
GY02	0.25	8.7 ± 0.3	$0.39^{+0.90}_{-0.32}$	$4.6^{+10.8}_{-3.8}$
D10	0.23	8.9 ± 0.3	$0.33^{+0.09}_{-0.08}$	$3.9^{+1.3}_{-1.1}$
S07	0.15	9.8 ± 0.3	$0.36^{+0.24}_{-0.16}$	$3.8^{+2.6}_{-1.8}$
D10	0.08	10.6 ± 0.3	$0.23^{+0.11}_{-0.08}$	$2.3^{+1.2}_{-0.9}$
M08	0.02	11.4 ± 0.3	$0.28^{+0.11}_{-0.08}$	$2.6^{+1.1}_{-0.9}$

Notes. (1) References for the different measurements: FM18; B12 – Barbary et al. (2012); GY02 – Gal-Yam, Maoz & Sharon (2002); S10 – Sharon et al. (2010); G08 – Graham et al. (2008); D10 – Dilday et al. (2010); S07 – Sharon et al. (2007); M08 – Mannucci et al. (2008). (2) Visibility-time-weighted average redshift of each cluster sample. (3) Delay time corresponding to a formation redshift between $z_f = 3-4$, given the visibility-time-weighted average redshift. (4) SN Ia rate per unit rest-frame B -band stellar luminosity, derived as explained in Section 2.3 for the FM18 sample, and as indicated in table 6 of FM18 for the previous measurements. (5) SN Ia rate per unit formed stellar mass, derived as explained in Section 2.3. The average B -band mass-to-light ratios from Pégase .3 models can be recovered by dividing columns (4) by (5).

a single star-formation burst, while the delay time range allowed for each cluster corresponds to the time elapsed between formation redshifts in the range $z_f = 3-4$ and the visibility-time-weighted average redshift of each cluster sample. In this uniform re-analysis or previous rates, we generally find higher mass-to-light ratios than previously assumed, and as a result the rates per unit mass for the low- z clusters are about 30 per cent lower than estimated in the original publications.

Table 3 summarizes the updated SN Ia rates based on the single-burst assumption, for the FM18 sample, for its two subsamples, and for the previous measurements compiled by Maoz & Graur (2017). For the previous measurements, the SN Ia rates per unit B -band luminosity are reproduced, unchanged, from table 6 of FM18, and the updated SN rates per unit formed stellar mass are based on the revised Pégase .3 mass-to-light ratios, as described above.

3 MODELLING

With the revised fluxes and SNe Ia numbers for the FM18 sample and the revised lower- z SNe Ia rates in hand, we now attempt to constrain the SN Ia DTD in cluster environments in the presence of extended star formation histories (SFHs) for the monitored FM18 clusters. To this end, we calculate SPS models, and find the range of SFH parameters that are consistent with the observed spectral energy distributions (SEDs) of the clusters in the *HST* bands. Simultaneously, we find the allowed range of the parameters describing the DTD, by fitting the observed number of SNe Ia in each cluster in FM18 to the predictions for a given SFH plus DTD combination, and further by

taking into account previous measurements of the SN Ia rate per unit formed stellar mass in lower- z cluster samples.

3.1 Star-formation parameters and SED fitting

We constrain the range of SFH parameters, that are consistent with the observed *HST* stellar fluxes for each cluster, by producing model SEDs with the `Pégase.3` SPS code. To investigate a broad range of possible SFHs for the clusters, yet without introducing too many free parameters, we assume, for each cluster, an exponentially decaying star formation rate (SFR),

$$\psi(t) = \frac{M_0}{\tau} \exp(-t/\tau), \quad (2)$$

having three parameters: (i) the initial star-formation redshift, z_f , which is equivalently represented with a parameter t , the cosmic time elapsed between z_f ($t = 0$) and z at which the cluster is observed; (ii) τ , the characteristic time of the exponential; and (iii) M_0 the asymptotically (at infinite time) formed stellar mass. The limit $\tau = 0$ corresponds to an instantaneous burst of star formation, with $\psi(t) = M_0 \delta(t)$ where δ is the Dirac delta function. In the SPS models, as before, we adopt an unevolving metallicity of $Z = 0.020$, based on the close-to-solar abundances generally observed in the massive early-type galaxies that make up the bulk of the stellar mass in galaxy clusters (e.g. Conroy et al. 2014), a Kroupa (2001) IMF between 0.1 and $100 M_\odot$, and Padova+AGB stellar evolution tracks. A fourth modelling parameter that we adopt, which relates the cluster's model SFH to its observed SED, is a rest-frame V -band dust extinction parameter A_V , based on the Cardelli, Clayton & Mathis (1989) extinction curve with reddening parameter $R_V = 3.1$, applied as a foreground dust screen to the synthesized galaxy light. Finally, we allow for a 'late' SFH component in the form of a second exponential burst, beginning 100 Myr before the cluster-observation epoch, with a fixed exponential time-scale of 100 Myr. This SFH model component opens up an option to account for the blue emission sometimes apparent in the cluster SEDs, and indeed in some individual cluster galaxies with indications of recent star formation. The choice of 100 Myr for both the delay and for the exponential time-scale of this component allows for a significant effect, both on the light (this late stellar population is still young and forming) and on the SN Ia rate (>40 Myr have elapsed, so the DTD has 'turned on', and the DTD is still near its peak SN Ia rate). This second burst of star formation is parametrized with $m = M'_0/M_0$, where M'_0 is the asymptotically formed stellar mass of the late burst. The late burst is assumed to have the same constant solar metallicity as the main burst, and to be unaffected by dust. The assumed SFH of each cluster is thus described by five parameters: t , τ , M_0 , A_V , and m , with M_0 being essentially a scaling factor that matches the calculated fluxes to the observed ones.

For each set of parameters, we use `Pégase.3` to derive a model spectrum for a unit stellar mass. For each cluster in the observed sample, we shift the model spectrum to the cluster redshift, and convert the model's luminosity density per unit wavelength interval, L_λ , to a flux density $f_\lambda = L_\lambda/[4\pi D_L^2(1+z)]$, where D_L is a cluster's luminosity distance, assuming standard cosmological parameters ($\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$). The spectrum is then folded through the *HST*+filter system bandpasses in F105W, F140W, and F160W, to obtain the model band-averaged flux in each observed band, which can be compared to the measured fluxes.

We use χ^2 as a figure of merit to gauge the agreement of an SFH model with the observed fluxes of a particular cluster:

$$\chi^2 = \sum_i \frac{(f_i - M_0 f_{i,\text{model}})^2}{df_i^2}, \quad (3)$$

where f_i is a measured integrated-galaxy-cluster flux in the i th band, df_i is its associated uncertainty (see Table 1), $f_{i,\text{model}}$ is the predicted flux per unit formed stellar mass for the given SFH model, and the best-fitting asymptotically formed stellar mass is

$$M_0 = \frac{\sum_i \frac{f_i f_{i,\text{model}}}{df_i^2}}{\sum_i \frac{f_{i,\text{model}}^2}{df_i^2}}, \quad (4)$$

i.e. the weighted average of the ratios $f_i/f_{i,\text{model}}$, given the observational uncertainties.

For each of the parameters of the SFH modelling, we explore the following ranges. The delay-time range corresponds to formation redshifts in the range $z_f = 3\text{--}4$ (see Section 1). We restrict the SFR time-scale τ to the range 0–1.2 Gyr. This is based on the finding by Salvador-Rusiñol et al. (2019) that <0.5 per cent of the stellar mass of massive early-type galaxies at $z \sim 0.4$ was formed in the preceding 2 Gyr. For a galaxy with an exponentially declining SFR, formed at $z_f = 4$, creation of <0.5 per cent of its mass between lookback times of 6 and 4 Gyr (i.e. $z \approx 0.4$ and 2 Gyr earlier), requires $\tau < 1.2$ Gyr. (A stricter constraint would result if formation were assumed at $z_f = 3$). This range in τ thus assures that our assumed SFHs for the clusters are consistent with the sensitive limits set by Salvador-Rusiñol et al. (2019) on the presence of any young stellar populations in early-type galaxies. We limit the amount of dust extinction to $A_V < 1.5 \text{ mag}$; the modelling of the individual SN Ia light curves in FM18, as well as SPS models fit to the spectra of the individual SN host galaxies in W20, both yield typical extinctions below this value. Early-type galaxies, as a rule, do not reveal evidence of much dust. The parameter m , embodying the effects of late-time star formation, is allowed to vary between 10^{-3} and 1.

To illustrate the SED fitting process, Fig. 1 shows the best-fitting spectrum and the range of allowed spectra for four of the FM18 clusters. While the parameter m is generally limited by the fits to small values between 10^{-3} and 10^{-2} , the other SFH parameters are weakly constrained by the cluster photometry alone, and thus they can span much of their allowed range.

3.2 DTD parameters and modelling of SN Ia numbers

In line with previous work, we will assume a DTD having a power-law form, with units of SN Ia rate per formed stellar mass,

$$\text{DTD}(t) = R_1 \left(\frac{t}{\text{Gyr}} \right)^\alpha, \quad (5)$$

where R_1 is the DTD rate at $t = 1$ Gyr. The SN Ia rate at time t in a galaxy cluster with a given SFH, $\psi(t)$, is the convolution of the SFH with the DTD,

$$R_{\text{Ia,model}}(t) = \int_{t_{\text{min}}}^t \psi(t-t') \times \text{DTD}(t') dt'. \quad (6)$$

As already noted, $t_{\text{min}} = 40$ Myr is the stellar evolution time until the appearance of the first white dwarfs in a stellar population. For the exponentially decaying SFH of equations (2) and (6) becomes

$$R_{\text{Ia,model}}(t) = \frac{M_0 R_1}{\tau} \int_{t_{\text{min}}}^t \left(\frac{t'}{\text{Gyr}} \right)^\alpha \exp\left(-\frac{t'-t}{\tau}\right) dt', \quad (7)$$

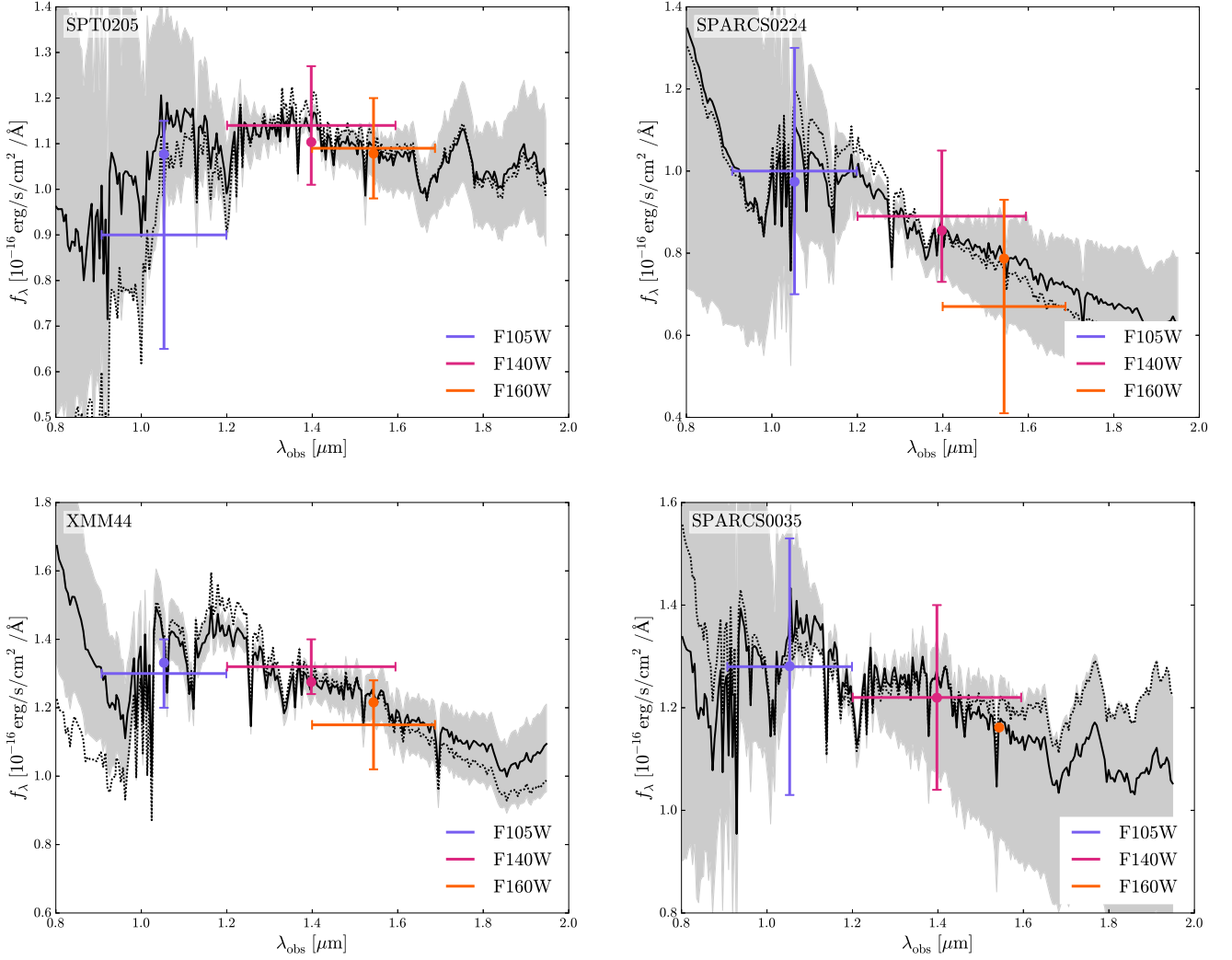


Figure 1. Observed fluxes and $\text{P}\acute{\text{e}}\text{gase.3}$ model spectra for four of the FM18 clusters. The dotted spectrum corresponds to the best-fitting (lowest- χ^2) model, the solid spectrum to the median of the allowed models, and the grey area to the 1σ range. Observed fluxes and uncertainties in the F105W, F140W, and F160W *HST* bands are indicated by vertical error bars. Horizontal error bars denote the filter band passes. Predicted fluxes, from folding the median model spectrum through the bandpass responses, are marked with filled circles.

which simplifies to $M_0 R_1 (t/\text{Gyr})^\alpha$ for an instantaneous burst. For our SFH models that include a second burst, 100 Myr before the redshift of observation, we add to the SN Ia rate from the main, initial burst, the SN Ia rate induced by this second burst.

The rest-frame monitoring times and the experiment’s detection-completeness factors for SNe Ia that reached maximum light before, during, and after each *HST* observing ‘season’ for each cluster, were derived based on simulations in FM18, and are reproduced in Table 1. The number of detected SNe Ia predicted for the experiment in a given cluster by a given SFH model and a particular DTD model is

$$N_{\text{Ia,model}} = R_{\text{Ia,model}}(t) \times \sum_i \tau_{\text{vis},i} \eta_i, \quad (8)$$

where the sum is both over the observing seasons (if more than one) and the periods before, during and after each season. For each cluster, the $\tau_{\text{vis},i}$ are the rest-frame monitoring times, i.e. the cluster rest-frame interval of actual monitoring, a $40/(1+z)$ -day interval before the monitoring period, and a $20/(1+z)$ -day interval after the monitoring period, and η_0 , η_{-40} , and η_{+20} are the corresponding detection-completeness factors. This expected number of SN Ia

can be compared to the actual number observed in each cluster, constraining both the SFH model and the DTD parameters.

Since the low- z cluster samples whose SN Ia rates were compiled in FM18 are observed at relatively large delay times, for which the assumption of passive evolution after a brief burst at high z is likely valid, their observed SN Ia rates per formed stellar mass (as revised in Section 2.3) are direct measures of the DTD at different delays. To obtain prior constraints on the DTD parameters imposed by these previous low- z SN Ia rate measurements, we quantify the agreement of a given DTD model (R_1, α) with the measurements through

$$\chi_{\text{low}}^2 = \sum_j \left(\frac{R_{\text{Ia},M,j} - \text{DTD}(t_j | R_1, \alpha)}{dR_{\text{Ia},M,j}} \right)^2, \quad (9)$$

where for the j th low- z cluster sample, t_j , $R_{\text{Ia},M,j}$, and $dR_{\text{Ia},M,j}$ are respectively the delay time, the rate per unit formed stellar mass, and its uncertainty, as indicated in Table 3. When $R_{\text{Ia},M,j} < \text{DTD}(t_j)$, we use the positive uncertainty; when $R_{\text{Ia},M,j} > \text{DTD}(t_j)$, the negative one. The prior distribution of R_1 and α , based on the low- z cluster SN Ia rates, is therefore the chi-square probability with one degree

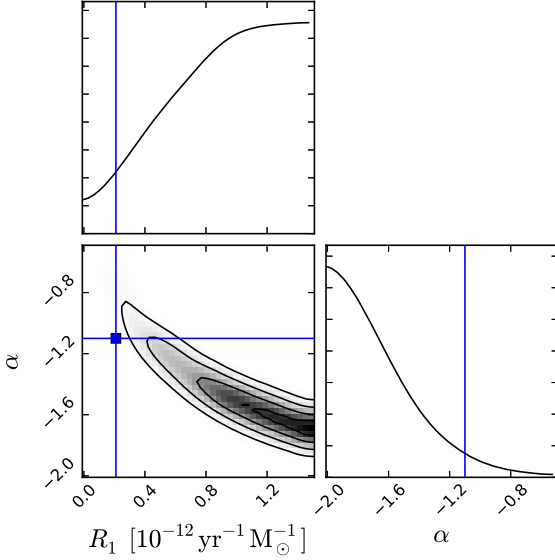


Figure 2. Prior distribution of the DTD parameters R_1 and α , based on the lower- z cluster sample SN rate measurements compiled by FM18 and revised in Section 2.3. The prior is assumed to follow a chi-square probability with one degree of freedom, as indicated in equation (10). For reference, the blue square highlighted with a solid cross corresponds to the field-galaxy DTD parameters from Maoz & Graur (2017): $R_1 = (0.21 \pm 0.02) \times 10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$ and $\alpha = -1.07 \pm 0.09$. Contours correspond to 0.5, 1, 1.5, and 2σ levels. The low- z cluster SN rates alone do not constrain strongly the DTD, with a strong anti-correlation between its two parameters.

of freedom

$$\mathcal{P}_{\text{low}}(R_1, \alpha) = \frac{\chi_{\text{low}}^{-1} e^{-\chi_{\text{low}}^2/2}}{2^{1/2} \Gamma(\frac{1}{2})}, \quad (10)$$

which is shown in Fig. 2. As noted in Maoz et al. (2010), where the SN Ia rates from these low- z clusters were already analysed, these low- z rates alone do not have sufficient leverage in terms of redshift range (translating to delay-time range) in order to individually constrain well the two DTD parameters, which are strongly anticorrelated in the fit. However, as we will see below, when combined, as a prior, with the high- z (short-delay) SN Ia measurements of FM18, the low- z measurements do significantly tighten the constraints on the DTD.

3.3 Simultaneous SFH and DTD model fits to the data

We use Bayesian inference to derive posterior probabilities for the SFH and DTD parameters, i.e.

$$\mathcal{P}(\theta|D) \propto \mathcal{P}(D|\theta) \times \mathcal{P}(\theta), \quad (11)$$

where θ indicates the parameters, D the data, $\mathcal{P}(\theta)$ the prior, and $\mathcal{P}(D|\theta)$ the likelihood. For describing the SFH of each of the 12 FM18 clusters, each with four SFH parameters (t , τ , m , and A_V), θ is a vector of $12 \times 4 + 2 = 50$ elements with the two parameters of the universal DTD (R_1 , α). D is a vector of $2 \times 32 + 12 + 2 \times 12 + 4 \times 18 = 172$ elements: 32 fluxes in different bands and their uncertainties, 12 redshifts, 12 counts of likely SNe Ia, 12 counts of possible SNe Ia, and 18 monitoring seasons and their detection-completeness factors.

The prior is assumed uniform for all SFH parameters, with the following boundaries: t corresponds to a formation redshift between $z_f = 3-4$, τ is between 0 (instantaneous burst) and 1.2 Gyr, A_V

between 0 and 1.5 mag, $\log m$ is between -3 and 0 (see last paragraph of Section 3.1, for a discussion of the boundaries chosen for τ and A_V). The DTD parameters α and R_1 are assumed to be, respectively, between -5 and 0 and between 0.01 and $5 \times 10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$. When considering only the FM18 clusters and their SNe Ia, we assume uniform priors for the DTD parameters. To take into account the constraints on the DTD parameters stemming from the low- z cluster measurements, we assume the prior $\mathcal{P}(\theta) = \mathcal{P}_{\text{low}}(R_1, \alpha)$ from equation (10).

The logarithm of the likelihood is the sum of probabilities

$$\log \mathcal{P}(D|\theta) = \sum_i \log \mathcal{P}_{\text{SFH},i} + \sum_i \log \mathcal{P}_{\text{FM18},i}, \quad (12)$$

where

$$\mathcal{P}_{\text{SFH},i} = \frac{\chi_i^{-1} e^{-\chi_i^2/2}}{2^{1/2} \Gamma(\frac{1}{2})} \quad (13)$$

is the chi-square probability with one degree of freedom associated with the χ^2 of equation (3) for cluster i , and

$$\mathcal{P}_{\text{FM18},i} = \frac{(N_{\text{Ia,model},i})^{N_{\text{Ia,obs},i}} e^{-N_{\text{Ia,model},i}}}{\Gamma(N_{\text{Ia,obs},i} + 1)} \quad (14)$$

is the Poisson probability of an observed SN Ia count $N_{\text{Ia,obs},i}$ in cluster i given a count predicted by a model with particular SFH and DTD parameters. The formulation of the Poisson probability with the gamma function Γ permits considering non-integer ‘observed’ counts of SNe – as in FM18, the observed count of SNe Ia in each of the clusters is taken as

$$N_{\text{Ia,obs}} = N_{\text{Ia}} + 0.5N_{\text{Ia}^?}, \quad (15)$$

where N_{Ia} and $N_{\text{Ia}^?}$ are, respectively, the number of likely and possible cluster SNe Ia, as indicated in Table 1. The two sums of equation (12) are over the 12 FM18 high- z clusters at the focus of this work.

To understand the results of the present analysis and any differences from previous studies of the cluster-environment SN Ia DTD, we consider four different setups, to gauge both the effect of assuming extended SFHs versus single instantaneous starbursts (as in previous studies) and the effect of considering the data for the lower- z clusters:

(i) **Setup 1.** Only the FM18 clusters and their SNe Ia are modelled, with their SFHs described by single short bursts. In this case, the model has 14 parameters: t for each of the 12 clusters, R_1 , and α . We assume $\tau = 0$, no extinction ($A_V = 0$), and no late burst ($m = 0$). Uniform priors are assumed for t , R_1 , and α .

(ii) **Setup 2.** Only the FM18 clusters, but allowing for extended SFHs in the clusters. In this case, the model has 50 parameters: t , τ , A_V , and m for each of the 12 clusters, R_1 , and α . Uniform priors are again assumed for all parameters.

(iii) **Setup 3.** The FM18 clusters are considered, together with low- z constraints on the two DTD parameters, but again with all SFHs described by single short bursts. As for setup 1, the model has 14 parameters: t for each of the 12 clusters, R_1 , and α . We assume $\tau = 0$, no extinction ($A_V = 0$), no second burst ($m = 0$). The prior for the DTD parameters, based on low- z cluster rate measurements, is $\mathcal{P}_{\text{low}}(R_1, \alpha)$ from equation (10). A uniform prior is assumed for t .

(iv) **Setup 4.** We model the FM18 clusters, permitting extended cluster SFHs, and considering the constraints on the two DTD parameters from the low- z cluster samples. As for setup 2, the model has 50 parameters: t , τ , A_V , and m for each of the 12 clusters, R_1 , and α . While uniform priors are assumed for the SFH parameters, the prior for the DTD parameters is $\mathcal{P}_{\text{low}}(R_1, \alpha)$ from equation (10). Amongst the four setups, this setup makes maximal use of the

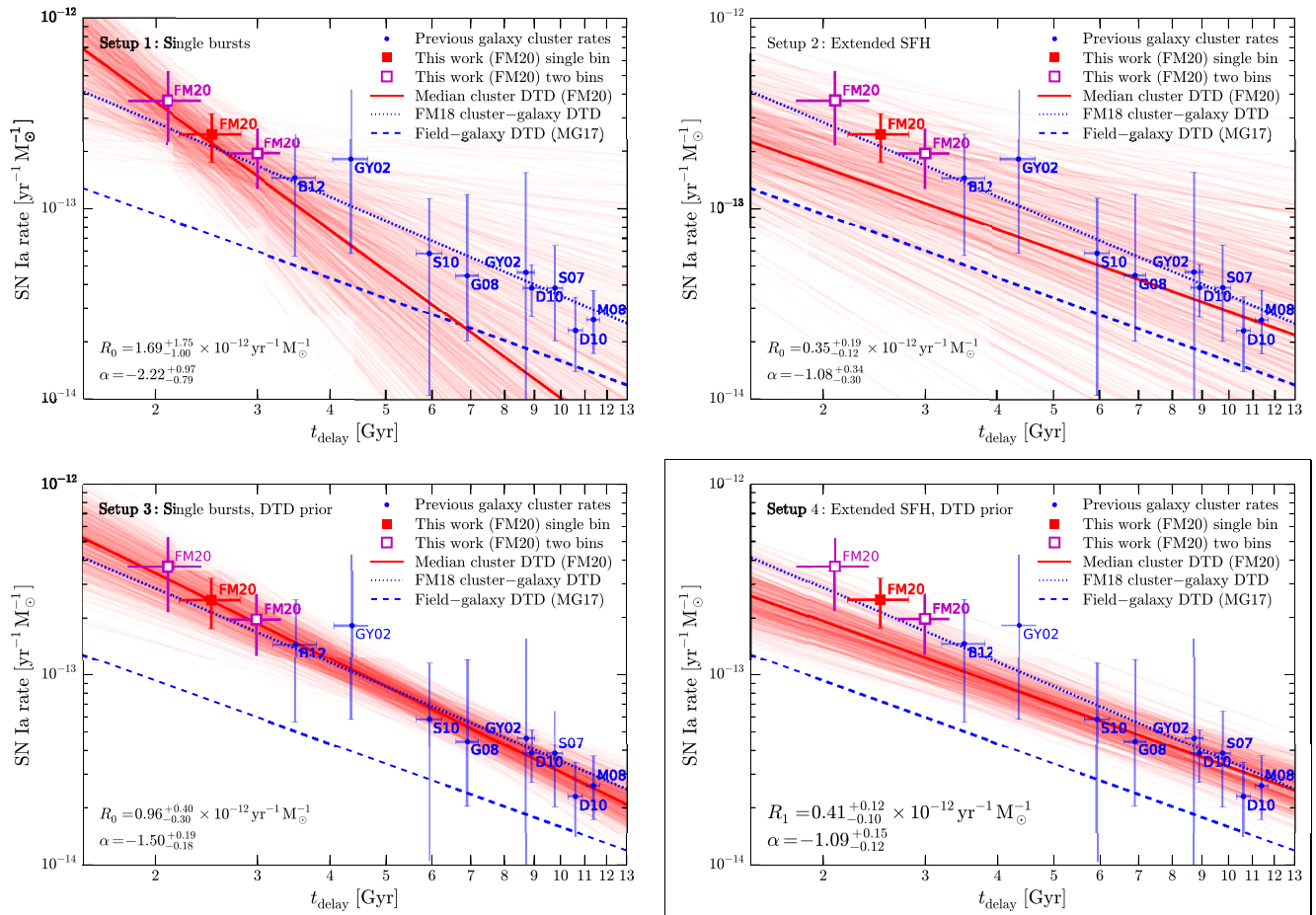


Figure 3. Cluster-environment power-law DTD models emerging from our MCMC Bayesian inference, given the four different modelling setups (see Section 3.3). Thin red lines sample the DTD parameters resulting from the MCMC, while the thick red line shows the DTD with the median parameters in the converged MCMC calculation. The cluster-environment best-fitting DTD from FM18 is plotted as a dotted blue line, and the field-galaxy DTD from Maoz & Graur (2017) as a dashed blue line. Data points, shown here for reference only (i.e. these data points are not being fit), are the SN Ia rates per unit formed stellar mass, in the FM18 cluster samples and in previously studied, lower- z cluster samples, all with values as revised in this paper and presented in Section 2.3. These plotted rates are based on the assumption that the stars in all the clusters were formed in a single burst between $z = 3$ and 4, an assumption that we forego in the model setups shown in the right-hand panels. Comparison of the two bottom panels shows that allowing, in the FM18 clusters, for extended SFHs (setup 4) that are consistent with the observed luminosities and colours, leads to a cluster SN Ia DTD (bottom right-hand panel) that is shallower and lower normalized than the DTD with the single-burst assumption (setup 3, bottom left-hand panel). The setup-4 cluster DTD, however, is still significantly higher normalized than the field-SN DTD, albeit only by a factor of 2–3.

observational data to constrain the DTD while allowing a wide range of SFHs. We base our conclusions on setup 4.

With the definition of the likelihood $\mathcal{P}(D|\theta)$ and the prior $\mathcal{P}(\theta)$, we construct the posterior probability $\mathcal{P}(\theta|D)$ using the Monte Carlo Markov Chain (MCMC) python implementation `emcee` by Foreman-Mackey et al. (2013). We have calculated the four different MCMC setups outlined above, to derive the posterior probabilities of the parameters describing the cluster SFHs and, more importantly, the DTD parameters. In the `emcee` implementation of the MCMC, we use 500 walkers. We monitor the convergence of the chains by following the walker positions, their medians and standard deviations, the evolution of the posterior probability, as well as the autocorrelation time. The integrated autocorrelation times are ~ 3000 iterations for setups 1 and 3 and $\gtrsim 10000$ for setups 2 and 4. We let the chains reach 100 000 iterations for the four setups, and estimate the posterior distributions over the last 10 000 iterations.

4 RESULTS

Fig. 3 shows the power-law DTDs resulting from the four different MCMC model-fitting setups outlined in Section 3.3; Fig. 4 gives, for each setup, the corresponding posterior distribution of the DTD parameters, R_1 and α . The upper panels show the DTDs resulting from fitting only the photometry and SN Ia numbers of the FM18 high-redshift cluster sample (setups 1 and 2), while ignoring constraints based on previous cluster SN Ia rate measurements. The lower panels, in contrast, do take into account as a prior the low- z constraints on the DTD parameters (setups 3 and 4). The two left-hand panels show the DTDs obtained when assuming single star-formation burst histories for all clusters, as in previous studies; the right-hand panels give the DTDs resulting when allowing, in the FM18 clusters, extended SFHs, the focus of this work. For reference purposes alone, in Fig. 3 we plot, beside the model DTDs, the SN Ia rates per unit formed stellar mass (assuming a single starburst) from Section 2.3, for the FM18 sample and for the low- z cluster

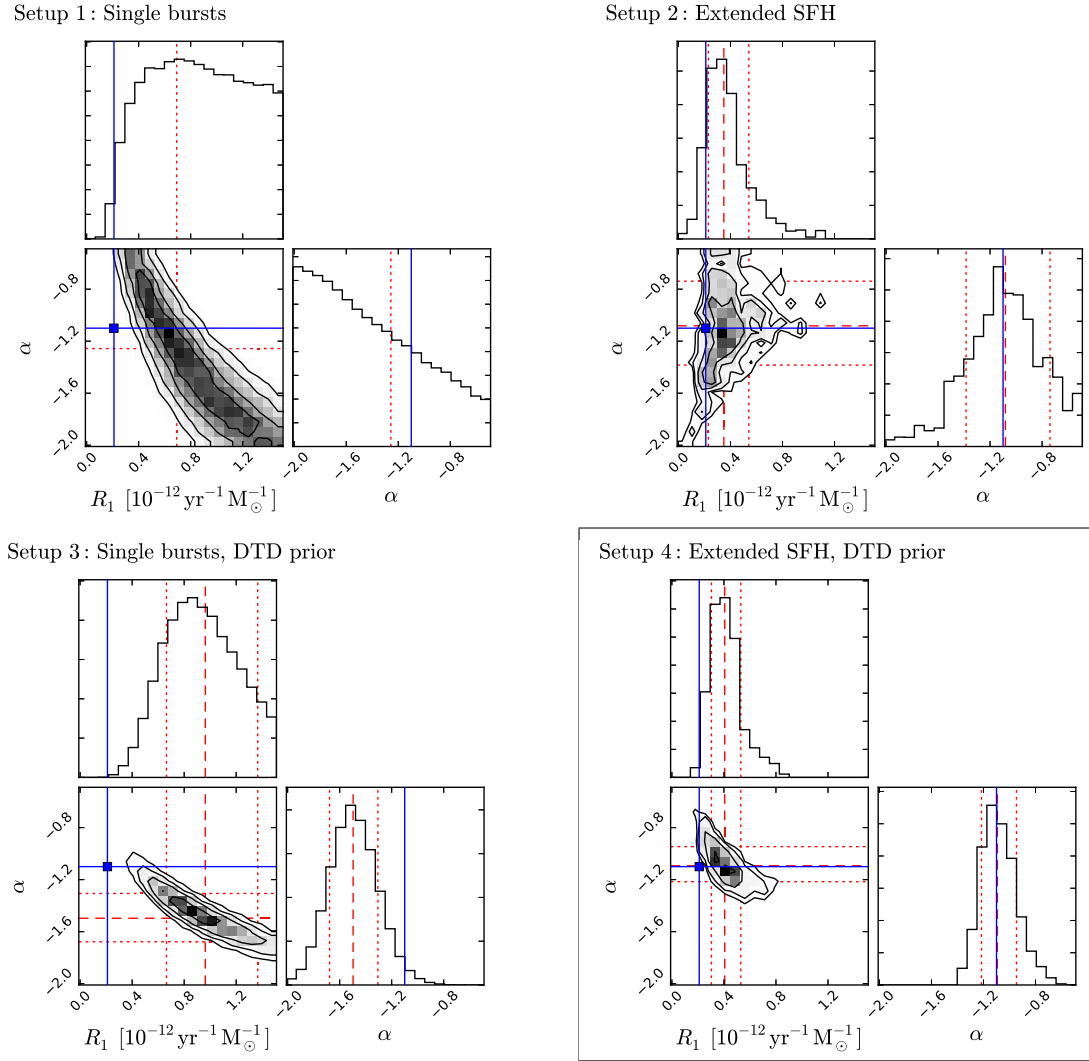


Figure 4. Posterior distributions of the DTD parameters R_1 and α for the four different setups outlined in Section 3.3, paralleling Fig. 4. The red dashed and dotted lines indicate the medians, and the $\pm 1\sigma$ 16th and 84th percentiles. The blue square highlighted with a solid cross corresponds to the field-galaxy DTD parameters from Maoz & Graur (2017). Contours correspond to 0.5, 1, 1.5, and 2σ levels. The upper panels illustrate how, when considering only the FM18 cluster SNe, enabling extended SFH significantly reduces the allowed range for the DTD parameters. The lower panels show that the addition also of priors based low-redshift cluster rate measurements further constrains the DTD parameters. In the lower right-hand panel, allowing both extended SFH and prior constraints leads to our main result: a DTD power-law slope α consistent with the one measured in field-galaxy environments, and with an amplitude significantly (3.8σ) higher, by a factor of 2–3, than the field-galaxy SN Ia DTD.

samples, all as revised in this paper (identical values are plotted for all of these rates in all four panels). We emphasize, however, that in none of these plots/setup are these sample- and sub-sample SN rates the observables that are being fit by the models. Rather, the observables being fit are the luminosities in the observed bands of each of the FM18 clusters, and the number of SNe Ia detected within each individual FM18 cluster. In the figures, the MCMC results are also compared to the cluster DTD previously obtained by FM18 and to the field-galaxy DTD from Maoz & Graur (2017).

The upper left-hand panels in Fig. 3 and in Fig. 4 show that, when considering only the FM18 clusters, with their limited range of short delay times around 2.5 Gyr, the assumption of single star-formation bursts (setup 1) leads to only weak constraints on the DTD parameters, with large covariance between the two values of the two DTD power-law parameters. This result largely echoes the results shown previously in Fig. 2, for a similar reason – a limited range

of short delays (previously long delays). Allowing extended SFHs for the FM18 clusters, however (setup 2, upper right-hand panels in Figs 3 and 4), significantly reduce both the allowed range for the DTD parameters and their correlation in the posterior distribution. This outcome already highlights the importance of considering extended SFHs for the high- z clusters with their corresponding short delay times. The SFH parameters of the walkers span most of the allowed ranges in setup 2 (except for the parameter m describing the second star formation burst, for which values above 0.1 are generally excluded). With an exponential star-formation-rate time-scale τ that can be as high 1.2 Gyr, the clusters are consistent with being observed within less than two star-formation e-folding times, effectively probing a larger range of DTD delay times. Remarkably, the DTD parameters obtained with setup 2, namely $R_1 = 0.35^{+0.19}_{-0.12} \times 10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$ and $\alpha = -1.08^{+0.34}_{-0.30}$, are already quite similar to (although more uncertain than) those we obtain from

the final setup 4, described below, that includes the leverage provided by the lower redshift cluster SN rates. This highlights a consistency between the DTD parameters indicated by the low- z and high- z cluster rates.

In the lower left-hand panels (setup 3) of Figs 3 and 4 one sees that the addition of the low- z SN rate data through the prior \mathcal{P}_{low} constrains better the DTD parameters, shrinking the allowed parameter space. More importantly, the setup 3 calculation is, in essence, a repeat of the analysis of FM18 (and other previous cluster SN Ia DTD analyses) that assumed single-burst cluster SFHs, only here with our revised high- z SN Ia numbers and cluster luminosities, and the revised low- z cluster SN Ia rates. Unsurprisingly, the best-fitting DTDs indeed go through the data points that are plotted for reference, data points that as noted, are based on the same single-burst assumption. With these revised numbers and the single-burst assumption, the DTD amplitude remains high as in FM18 (i.e. several times higher than the field-DTD values), and the DTD power-law index, $\alpha = -1.50^{+0.19}_{-0.18}$, is even steeper than its $\alpha = -1.30^{+0.23}_{-0.16}$ value in FM18, which in turn is steeper than the field DTD value. The results of setup 3 thus confirm and strengthen the conclusions of FM18, of a high-normalized and steep cluster-environment DTD, when assuming single-burst cluster SFHs.

Finally, in the lower right-hand panels of Figs 3 and 4 (setup 4), we present the main results of this work – the cluster-environment DTD as constrained by the FM18-sample cluster and SN Ia data, and by the prior lower- z cluster-sample SN Ia rates, while allowing, for the FM18 clusters, extended SFHs that are consistent with the data. We obtain $R_1 = 0.41^{+0.12}_{-0.10} \times 10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$ and $\alpha = -1.09^{+0.15}_{-0.12}$. In terms of Hubble-time-integrated SN Ia number per formed stellar mass, this corresponds to $N/M_{\text{star}} = 3.1^{+1.1}_{-1.0} \times 10^{-3} M_{\odot}^{-1}$. On the one hand, these parameter values are within 1σ of the FM18 values, thus confirming with a different methodology, with revised data and with more realistic SFH assumptions, the FM18 conclusion that the SN Ia DTD in cluster environments is higher normalized than in field galaxies. The field-galaxy DTD parameters are $R_1 = 0.21 \pm 0.02 \times 10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$, $\alpha = -1.07 \pm 0.09$, and $N/M_{\text{star}} = 1.3 \pm 0.1 \times 10^{-3} M_{\odot}^{-1}$ (Maoz & Graur 2017). On the other hand, the dichotomy in DTD normalization between the two environments is decreased compared to the results in FM18, and is now only at the level of a factor of ~ 2 – 3 . Furthermore, the DTD power-law index α now appears to be fully consistent, at $\alpha \approx -1.1$, in both cluster and field environments, rather than being steeper in clusters. From comparison of the setup 3 and setup 4 results, it is clear that this convergence in DTD parameters comes about from the more-physical assumption of extended cluster SFHs that we have adopted in this work. Nevertheless the difference between the cluster and field DTD parameters remains significant. From the difference in log likelihood between the best-fitting (setup 4) cluster DTD solution and that for the parameters of the field-galaxy DTD, we deduce a 3.8σ significance.

We summarize in Table 4 the DTD parameters obtained with MCMC setups 3 and 4, and those from FM18 and Maoz & Graur (2017). In Table 5, we list the SFH parameters for the individual FM18 clusters from the MCMC fit for the adopted setup 4 calculation. For each cluster, we list also its model fluxes (which can be compared to Table 1) and the predicted versus observed SN Ia numbers.

5 CONCLUSIONS

We have re-analysed the data resulting from the near-infrared *HST* imaging observations, described and analysed in FM18, of a sample of galaxy clusters at redshifts $z = 1.13$ – 1.75 that were monitored

Table 4. SN Ia DTD parameters.

Setup (1)	α (2)	R_1 [$10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$] (3)	N/M_{star} ($10^{-3} M_{\odot}^{-1}$) (4)
Cluster-environment DTD (this work)			
Setup 3	$-1.50^{+0.19}_{-0.18}$	$0.96^{+0.40}_{-0.30}$	$11.3^{+6.8}_{-5.0}$
Setup 4	$-1.09^{+0.15}_{-0.12}$	$0.41^{+0.12}_{-0.10}$	$3.1^{+1.1}_{-1.0}$
Cluster-environment DTD (FM18)			
–	$-1.30^{+0.23}_{-0.16}$	$0.7^{+0.1}_{-0.2}$	$5.4^{+2.3}_{-2.3}$
Field-galaxy DTD (Maoz & Graur 2017)			
–	-1.07 ± 0.09	0.21 ± 0.02	1.3 ± 0.1

Notes. (1) MCMC calculation setup; setup 3 assumes single-burst star-formation histories (SFHs), as in previous cluster DTD studies; setup 4, which leads to our final result, allows for more realistic, extended, SFHs in the higher redshift FM18 clusters. (2) DTD power-law index α , as expressed in equation (5). (3) DTD normalization R_1 , which is the SN Ia rate per unit formed stellar mass at 1 Gyr delay time. (4) Hubble-time-integrated SN Ia number per formed stellar mass.

for SNe. To re-cap, FM18 discovered the transient events in these data, selected those events consistent with being SNe Ia hosted by the clusters, and performed multiband galaxy-light photometry for every cluster in the sample. The photometry, combined with single-burst stellar population synthesis models, was used by FM18 to estimate the initially formed stellar mass of each cluster. Combining the detected SN Ia numbers, the formed stellar masses, the monitoring period lengths, and the detection efficiencies of the experiment, FM18 derived the SN Ia rate per unit formed mass, for this sample. Finally, FM18 fit power-law SN Ia DTDs to the new SN Ia rates, combined with the SN Ia rates measured for previous, lower redshift, cluster samples. Confirming previous results (Maoz et al. 2010; Maoz & Graur 2017), FM18 obtained a cluster-environment DTD that is slightly steeper ($\alpha \approx -1.3$) than seen in field-galaxy environments ($\alpha \approx -1.1$) but, more important, considerably higher normalized (i.e. with a larger amplitude R_1) than the field-galaxy DTD. The high normalization, if real, would indicate a high time-integrated efficiency of SN Ia production in galaxy clusters. While this result was intriguing, concern remained that it was affected by the single-starburst assumption that it involved. For the high- z FM18 clusters, observed only a few rest-frame Gyr after initial star formation at $z = 3$ – 4 , some ongoing star formation in some of the cluster galaxies is possible and, in fact, evident to some degree directly in the observations (e.g. blue galaxy colours, evidence of dust). Not accounting properly for such star formation may have led, in the FM18 analysis, to overestimation of the stellar masses, an overestimate of the delay time probed at a given redshift, or both. This, in turn, could potentially influence the derived DTD.

In this paper, we have re-visited the problem, but now relaxing the single-burst assumption for the FM18 clusters, and instead allowing, for each cluster, extended SFHs that are consistent with its observed *HST* photometry. In the said photometry, we realized that FM18 had sometimes inadvertently over-subtracted the light of galaxies that are foreground to the clusters. We therefore re-measured all of the photometry from the *HST* images, and have reported the revised values, which are somewhat higher than in FM18. Furthermore, considering the results of follow-up spectroscopy of some of the SN host galaxies, as reported by Williams et al. (2020), we have revised (slightly upwards) the numbers of SNe Ia hosted by a few of the clusters. Our revised SN Ia rates per unit formed stellar mass (still under the single-burst assumption), based on the revised photometry and SN Ia numbers, are only mildly different from the rates in FM18.

Table 5. SFH parameters and SNe Ia numbers resulting from setup 4 (with extended SFHs and a DTD prior).

Cluster name (1)	t (2)	τ (3)	A_V (4)	$\log(m)$ (5)	M_0 (6)	$f_{F105W,model}$ (7)	$f_{F140W,model}$ (8)	$f_{F160W,model}$ (9)	$N_{Ia,model}$ (10)	$N_{Ia,obs}$ (11)
IDCS1426	$1.76^{+0.20}_{-0.19}$	$0.6^{+0.4}_{-0.4}$	$0.73^{+0.45}_{-0.45}$	$-2.36^{+0.90}_{-1.14}$	$0.81^{+0.35}_{-0.35}$	$0.72^{+0.24}_{-0.24}$	$0.63^{+0.02}_{-0.02}$	$0.57^{+0.05}_{-0.05}$	$1.6^{+0.9}_{-0.6}$	1.5
ISCS1432	$2.60^{+0.19}_{-0.20}$	$0.5^{+0.4}_{-0.3}$	$0.57^{+0.46}_{-0.36}$	$-2.56^{+0.37}_{-0.57}$	$1.22^{+0.29}_{-0.29}$	$0.88^{+0.10}_{-0.10}$	$0.92^{+0.01}_{-0.01}$	$0.90^{+0.04}_{-0.04}$	$1.7^{+0.9}_{-0.5}$	1.0
MOO1014	$2.95^{+0.19}_{-0.19}$	$0.7^{+0.4}_{-0.4}$	$0.58^{+0.55}_{-0.41}$	$-1.91^{+0.23}_{-0.32}$	$1.37^{+0.32}_{-0.32}$	$2.16^{+0.15}_{-0.15}$	$1.67^{+0.08}_{-0.08}$	$1.51^{+0.11}_{-0.11}$	$2.3^{+1.1}_{-0.7}$	2.0
MOO1142	$3.32^{+0.18}_{-0.23}$	$0.6^{+0.3}_{-0.4}$	$0.60^{+0.55}_{-0.40}$	$-2.20^{+0.08}_{-0.33}$	$1.99^{+0.46}_{-0.46}$	$2.31^{+0.00}_{-0.00}$	$2.11^{+0.01}_{-0.01}$	$2.01^{+0.05}_{-0.05}$	$1.2^{+0.5}_{-0.4}$	0.0
SPARCS0224	$2.11^{+0.20}_{-0.21}$	$0.7^{+0.4}_{-0.4}$	$0.65^{+0.51}_{-0.37}$	$-2.07^{+0.46}_{-0.90}$	$1.07^{+0.43}_{-0.43}$	$0.98^{+0.16}_{-0.16}$	$0.85^{+0.07}_{-0.07}$	$0.78^{+0.11}_{-0.11}$	$0.8^{+0.5}_{-0.3}$	1.5
SPARCS0330	$2.12^{+0.18}_{-0.19}$	$0.7^{+0.3}_{-0.3}$	$0.58^{+0.50}_{-0.43}$	$-1.82^{+0.52}_{-0.76}$	$0.74^{+0.32}_{-0.32}$	$0.98^{+0.16}_{-0.16}$	$0.75^{+0.09}_{-0.09}$	$0.66^{+0.13}_{-0.13}$	$1.0^{+0.7}_{-0.4}$	0.5
SPARCS1049	$1.88^{+0.20}_{-0.19}$	$0.8^{+0.3}_{-0.4}$	$0.77^{+0.46}_{-0.40}$	$-1.67^{+0.15}_{-0.56}$	$1.01^{+0.35}_{-0.35}$	$1.30^{+0.12}_{-0.12}$	$0.93^{+0.02}_{-0.02}$	$0.80^{+0.01}_{-0.01}$	$0.5^{+0.4}_{-0.2}$	0.0
SPARCS0035	$2.81^{+0.19}_{-0.18}$	$0.7^{+0.3}_{-0.4}$	$0.61^{+0.42}_{-0.36}$	$-2.32^{+0.17}_{-0.87}$	$1.34^{+0.30}_{-0.30}$	$1.28^{+0.00}_{-0.00}$	$1.22^{+0.03}_{-0.03}$	$1.16^{+0.03}_{-0.03}$	$1.4^{+0.7}_{-0.4}$	0.5
SPT0205	$2.84^{+0.21}_{-0.16}$	$0.5^{+0.4}_{-0.3}$	$0.75^{+0.43}_{-0.43}$	$-2.73^{+0.47}_{-0.63}$	$1.61^{+0.40}_{-0.40}$	$1.00^{+0.12}_{-0.12}$	$1.11^{+0.01}_{-0.01}$	$1.09^{+0.04}_{-0.04}$	$2.9^{+1.4}_{-0.9}$	4.5
SPT2040	$2.39^{+0.17}_{-0.19}$	$0.7^{+0.3}_{-0.4}$	$0.69^{+0.43}_{-0.41}$	$-2.45^{+0.42}_{-0.68}$	$2.31^{+0.77}_{-0.77}$	$1.76^{+0.27}_{-0.27}$	$1.69^{+0.08}_{-0.08}$	$1.62^{+0.15}_{-0.15}$	$3.3^{+1.5}_{-1.0}$	3.0
SPT2106	$3.46^{+0.14}_{-0.17}$	$0.7^{+0.4}_{-0.4}$	$0.67^{+0.52}_{-0.50}$	$-1.32^{+0.12}_{-0.15}$	$1.04^{+0.32}_{-0.32}$	$3.55^{+0.04}_{-0.04}$	$1.98^{+0.00}_{-0.00}$	$1.67^{+0.03}_{-0.03}$	$1.9^{+1.1}_{-0.6}$	1.0
XMM44	$2.17^{+0.21}_{-0.17}$	$0.7^{+0.3}_{-0.4}$	$0.72^{+0.43}_{-0.44}$	$-2.37^{+0.29}_{-0.62}$	$1.81^{+0.50}_{-0.50}$	$1.31^{+0.07}_{-0.07}$	$1.28^{+0.03}_{-0.03}$	$1.22^{+0.04}_{-0.04}$	$1.3^{+0.6}_{-0.5}$	1.0

Notes. (1) Cluster ID; (2) Delay time, in Gyr; (3) Exponential timescale, in Gyr; (4) Visible extinction, in mag; (5) Relative mass m of the second burst, with respect to the first burst; (6) Best-fitting asymptotically formed stellar mass, in $10^{13} M_{\odot}$; (7)–(9) Total modelled galaxy cluster flux in the F105W, F140W, and F160W bands, in $10^{-16} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$; (10) Modeled SN Ia count; (11) Observed SN Ia count (cf. equation (15)).

Finally, we have also reviewed the SN Ia rates per unit mass in previous, lower- z cluster samples analysed by FM18. Consistently using single-burst SPS models and the mass-to-light ratios that they predict, we have revised the low- z cluster SN Ia rates (typically downward by ~ 30 per cent).

To constrain the cluster DTD while allowing for extended SFHs, we have fit, using an MCMC process, the observed photometry to the SEDs of simplified but broad families of parametrized SFHs, as predicted by SPS models (including also dust extinction). Simultaneously, parametrized power-law DTDs were fit to the data, by comparing the number of SNe Ia predicted in a cluster, given its model-assumed SFH, to the actual number observed in that cluster. In contrast to FM18, where the total number of SNe Ia in the cluster sample was considered, and the sample’s total stellar mass, here we avoid this loss of information due to averaging, by comparing in each cluster, with its SFH and mass constrained by the photometry, the predicted and the observed SN Ia numbers. The treatment using the individual clusters, which span a range of redshifts and hence of delay times, provides further leverage on the DTD. In the MCMC calculation, to account for the measured SN Ia rates in lower- z cluster samples, we have used them as a prior on the DTD parameters. The SN Ia rate data that are being fit have been slightly revised, as described above, and so in order to understand the cause behind any change in the results compared to those of FM18, we have performed the MCMC fits using several different setups. In particular, in one of the setups (No. 3) we have fit for the DTD with the revised data and our new methodology (i.e. fitting the SN Ia numbers in individual clusters, using the low- z cluster SN Ia rates as a prior), but limiting the SFHs of all clusters to a single brief starburst, as in FM18 and previous studies. In MCMC setup No. 4, which yields our new results, we fit the same data, but now allowing for extended SFHs in the FM18 clusters, as described above.

The DTD emerging from our setup 3 is similar in amplitude to the DTD of FM18 (compare, e.g. their height at an intermediate delay of ~ 5 Gyr, in fig. 3, bottom left-hand panel). Our new DTD for this setup is, however, slightly steeper than the one in FM18 ($\alpha = -1.50 \pm 0.19$, versus $\alpha = -1.30^{+0.23}_{-0.16}$, respectively). This is presumably a result of the decrease in our revised low- z SN Ia rates, which ‘pull’ the DTD down at long delays. This shows that the revised data numbers we

have used here do not lead to a DTD notably different from that of FM18; it is just as high-normalized, compared to the field-galaxy DTD, and if anything, has even a steeper slope, α , compared to the field DTD found by FM18.

Proceeding to MCMC setup 4 (fig. 3, bottom right-hand panel) fitting all of the data while allowing extended SFHs in the FM18 clusters, we have found that the DTD slope becomes shallower, $\alpha = -1.09^{+0.15}_{-0.12}$, and is now indistinguishable from the slope of the field DTD, $\alpha \approx -1.1$. The amplitude of the cluster DTD has become lower and better constrained than in FM18, $R_1 = 0.41^{+0.12}_{-0.10} \times 10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$ (which translates to a Hubble-time-integrated SN Ia number per formed stellar mass, $N/M_{\text{star}} = 3.1^{+1.1}_{-1.0} \times 10^{-3} M_{\odot}^{-1}$), compared to $R_1 = 0.96^{+0.40}_{-0.30} \times 10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$ in the single-burst setup 3. This difference in the resulting DTD shows that, indeed, accounting for extended SFHs in high- z clusters is important for reliably recovering the DTD in cluster environments: the DTD becomes lower than under the single-burst assumption, and flatter – as flat as the field DTD. However, our new cluster DTD is still significantly higher normalized, by a factor of 2–3, compared to the field DTD [$R_1 = (0.21 \pm 0.02) \times 10^{-12} \text{ yr}^{-1} M_{\odot}^{-1}$, or $N/M_{\text{star}} = (1.3 \pm 0.1) \times 10^{-3} M_{\odot}^{-1}$], even if the dichotomy between cluster and field DTDs has narrowed.

From a purely random-statistical perspective, the difference that we have now confirmed, between the cluster- and field-DTD normalizations, is significant, and therefore begs physical explanations (see Section 1 for some possibilities). However, we should also remember that the measurement of SN rates is fraught with systematic uncertainties. We have made an effort to estimate those systematics that we know of, and to include them in our error budgets, e.g. the uncertainty in SN Ia numbers because of uncertainty in transient classification and in host galaxy cluster membership; or the uncertainty in cluster luminosity because of the uncertainty in the choice of the image regions that are representative for the subtraction of foreground and background light contributions. None the less, those estimates of systematic uncertainty could be wrong. Additional systematics that we are unaware of could also be in play, perhaps related to the fact that all of the measurements in the process, from SN detection to the characterization of stellar populations in galaxies, pertain to different types of galaxies – massive early-types in clusters, versus late-types

in the field. Considering all of the above, we believe that the current evidence for a cluster DTD that is normalized high by a factor of 2–3, compared to the field-galaxy DTD, remains strong and formally significant, but also that caution and further study are warranted.

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DATA AVAILABILITY

The *HST* imaging data used in this work are available through the data archive at the Space Telescope Science Institute. The processed data from this article are available in the article and its tables. The codes used for the MCMC modelling are available at https://github.com/JonathanFreundlich/DTD_MCMC.

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