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Ultraviolet spectra of extreme nearby star-forming regions: Evidence for an overabundance of very massive stars

Peter Senchyna¹,^{1,2}★† Daniel P. Stark,¹ Stéphane Charlot,³ Jacopo Chevallard³,³ Gustavo Bruzual⁴ and Alba Vidal-García⁵

¹Steward Observatory, University of Arizona, 933 N Cherry Ave, Tucson, AZ 85721, USA

²Carnegie Observatories, 813 Santa Barbara Street, Pasadena, CA 91101, USA

³Sorbonne Université, CNRS, UMR7095, Institut d'Astrophysique de Paris, F-75014 Paris, France

⁴Instituto de Radioastronomía y Astrofísica, UNAM, Campus Morelia, Michoacan, México C.P. 58089, México

⁵Laboratoire de Physique de l'Ecole normale supérieure, ENS, Université PSL, CNRS, Sorbonne Université, Université de Paris, F-75005 Paris, France

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ABSTRACT

As deep spectroscopic campaigns extend to higher redshifts and lower stellar masses, the interpretation of galaxy spectra depends increasingly upon models for very young stellar populations. Here we present new *HST*/COS ultraviolet spectroscopy of seven nearby (<120 Mpc) star-forming regions hosting very young stellar populations (~4–20 Myr) with optical Wolf–Rayet stellar wind signatures, ideal laboratories in which to benchmark these stellar models. We detect nebular C III] in all seven, but at equivalent widths uniformly <10 Å. This suggests that even for very young stellar populations, the highest equivalent width C III] emission at ≥ 15 Å is reserved for inefficiently cooled gas at metallicities at or below that of the SMC. The spectra also reveal strong C IV P-Cygni profiles and broad He II emission formed in the winds of massive stars, including some of the most prominent He II stellar wind lines ever detected in integrated spectra. We find that the latest stellar population synthesis prescriptions with improved treatment of massive stars nearly reproduce the entire range of stellar He II wind strengths observed here. However, we find that these models cannot simultaneously match the strongest wind features alongside the optical nebular line constraints. This discrepancy can be naturally explained by an overabundance of very massive stars produced by a high incidence of binary mass transfer and mergers occurring on short $\lesssim 10$ Myr time-scales, suggesting these processes may be crucial for understanding systems dominated by young stars both nearby and in the early Universe.

Key words: stars: massive – galaxies: evolution – galaxies: stellar content – ultraviolet: galaxies.

1 INTRODUCTION

Nebular line emission represents one of the most accessible windows on star formation in galaxies at cosmological redshifts. The measurement of rest UV–optical emission lines via near-infrared spectroscopy with the *James Webb Space Telescope* (*JWST*) and the first extremely large ground-based telescopes (ELTs) promises to unveil the chemical abundances and star formation histories of large samples of galaxies into the reionization era at $z > 6$ in the coming years (Stark 2016; Kewley, Nicholls & Sutherland 2019), approaching some of the earliest epochs of star formation in the Universe. Successful extraction of constraints on physical parameters of interest from these spectra where dominated by stellar photoionization will rely on models for the ionizing radiation output of populations of stars as a function of metallicity and age (stellar population synthesis) coupled to radiative transfer codes to predict the emergent nebular spectrum this radiation powers (photoionization modelling). While many sophisticated stellar population synthesis models incorporating stellar physics canonically neglected (such as

rotation and binary mass transfer; e.g. Leitherer et al. 2014; Eldridge et al. 2017; Götzberg et al. 2019) and new frameworks for performing inference using these predictions (e.g. Chevallard & Charlot 2016; Leja et al. 2017) have emerged in recent years, these predictions remain largely untested under the extreme star formation conditions we expect to encounter routinely in reionization era galaxies.

Broad-band SED fitting has provided an early glimpse of the nebular properties of star-forming systems at $z > 6$, revealing a population significantly different from typical galaxies at lower redshifts. Leveraging the fact that the strong rest-optical lines contaminate the *Spitzer*/IRAC [3.6] and [4.5] micron bands at $z \sim 6$ –9, several authors have robustly demonstrated that the specific star formation rates (sSFRs) of typical systems at these redshifts are well in-excess of galaxies at lower redshift (Labbé et al. 2013; Smit et al. 2014). Indeed, Endsley et al. (2021) have recently shown that 20 per cent of the most luminous systems at $z \sim 7$ reveal evidence for extremely high equivalent width optical nebular emission corresponding to sSFRs $\gtrsim 50 \text{ Gyr}^{-1}$, indicative of a substantial population of galaxies dominated entirely by massive stars formed within the last <20 Myr. Accurately interpreting the nebular line emission detected in such extreme systems will rely sensitively on the accuracy of models for the ionizing radiation fields powered by very young populations of massive stars.

* E-mail: psenchyna@carnegiescience.edu

† Carnegie Fellow.

Deep pilot near-infrared spectroscopy with large aperture ground-based telescopes has provided an early preview of the spectra that *JWST* will make commonplace. These data have revealed surprisingly prominent emission from doubly- and triply-ionized carbon in the rest-UV of several intrinsically luminous or gravitationally lensed systems at $z > 6$ (Stark et al. 2015a,b; Mainali et al. 2017; Schmidt et al. 2017; Hutchison et al. 2019). While emission in C III] $\lambda\lambda 1907, 1909$ (hereafter C III]) at such high equivalent widths ($\gtrsim 15 \text{ \AA}$) is rare in typical massive galaxies at lower redshifts (Du et al. 2017), a growing body of work at $z \sim 2$ with VIMOS and MUSE and in the local Universe ($\lesssim 100 \text{ Mpc}$) with the Cosmic Origins Spectrograph onboard the *Hubble* Space Telescope (*HST*/COS) has revealed that this semiforbidden doublet is routinely detected in subsolar metallicity galaxies at very high sSFR (Rigby et al. 2015; Berg et al. 2016; Maseda et al. 2017; Peña-Guerrero et al. 2017; Senchyna et al. 2017; Nakajima et al. 2018; Du et al. 2020). However, even the latest stellar population synthesis models struggle to reproduce the highest equivalent width emission in C III] detected at $z \sim 2\text{--}6$. Though AGN may contribute to some of this extreme emission especially at intermediate redshifts (Le Fèvre et al. 2019), the multiple detections of C III] in-excess of 15 \AA in $z > 6$ Lyman-break galaxies may also be the product of very prominent metal-poor massive stellar populations for which theoretical ionizing spectra predictions are not fully calibrated and plausibly inaccurate.

These results indicate that our ability to correctly interpret and extract constraints on underlying physics from nebular line emission will depend fundamentally on models for metal-poor stellar populations at very young ages. The ionizing spectrum powered by massive stars is essentially entirely theoretical at all metallicities and ages since the earliest stars observed directly at the extreme ultraviolet (EUV) energies relevant for H II region photoionization are B giants (Bowyer, Drake & Vennes 2000). Our best hope of calibrating these models is through analysis of nearby star-forming regions that can be studied in great detail. Much attention has traditionally been paid to the ability of population synthesis models to reproduce the distribution of strong-line ratios for typical star-forming systems, an important litmus test for their veracity (e.g. McGaugh 1991; Bresolin, Kennicutt & Garnett 1999; Charlot & Longhetti 2001; Byler et al. 2017; Xiao, Stanway & Eldridge 2018). However, many of the uncertain effects of stellar modelling (binary mass transfer, fast rotation, IMF variation) are degenerate in this space, especially at very young ages where constraints from large spectroscopic surveys are relatively lacking. More crucially, this does not test the core assumption underlying nebular emission line modelling: that model fits to the nebular emission lines alone can accurately constrain the properties of the responsible stellar population such as metallicity and age.

The difficulty in reproducing the first rest-UV nebular line detections at $z > 6$ underscores the urgent need for model testing for very young stellar populations. Modern stellar population synthesis results have diverged significantly in the past decade, with discrete improvements to the treatment of stellar winds and atmospheres, binary mass transfer, and rotation yielding predictions that diverge substantially at fixed physical parameters, especially for hard ionizing radiation (e.g. Wofford et al. 2016). If models predict the wrong stellar content and ionizing radiation field for galaxies dominated by a burst in star formation at the very young ages and subsolar metallicities that many reionization-era systems reside at, physical inferences including on star formation histories and metallicities from even the highest quality *JWST* spectra will suffer from potentially dramatic systematic uncertainties. Fortunately, spectroscopy of relatively bright star-forming regions in the local Universe can unveil far more detail

than just nebular line strengths. In particular, ultraviolet and deep optical observations especially at moderately subsolar metallicities $Z/Z_{\odot} \gtrsim 0.2$ directly accesses continuum signatures of the massive stars producing the ionizing radiation and nebular emission, including very prominent wind features formed in the dense radiatively driven outflows these stars power (e.g. Allen, Wright & Goss 1976; Brinchmann, Kunth & Durret 2008b; Leitherer et al. 2011; Wofford et al. 2014). With direct information about the massive stars present in these star-forming regions, far more stringent tests of stellar population models can be conducted.

In this paper, we present new *HST*/COS UV spectroscopy for seven local ($\lesssim 100 \text{ Mpc}$) star-forming regions specifically selected to enable such a test of theoretical prescriptions for young stellar populations. In particular, we focus on the subset of objects from the Shirazi & Brinchmann (2012) sample of star-forming He II emitters in SDSS with evidence for broad Wolf-Rayet (WR; see e.g. Crowther 2007) stellar wind emission at He II $\lambda 4686$, produced by some combination of very massive and envelope-stripped or spun-up stars (Section 2). These systems have properties similar by selection to the two prominent C III]-emitters at $12 + \log O/H \simeq 8.3$ presented in Senchyna et al. (2017, hereafter S17), with optical nebular emission and inferred ages comparable to that inferred for the highest sSFR systems in the reionization era (Section 3). Besides ensuring that direct constraints can be placed on the stellar populations present, the presence of prominent WR wind emission evinces a significant population of very hot stars. The UV data provide access both to nebular emission in C III] and O III] $\lambda\lambda 1661, 1666$, as well as the He II $\lambda 1640$ recombination line and C IV $\lambda\lambda 1548, 1550$ resonant doublet. At these metallicities and ages, both C IV and He II will be dominated by strong stellar wind features (a P-Cygni feature dominated by O-stars, and broad emission associated with very massive and WR stars, respectively; Section 4). These stellar wind features underlying the strong nebular emission lines enable a direct test of stellar population synthesis prescriptions (Section 5). The degree to which these models are able to match both the wind features and the nebular emission powered by the massive stars provides crucial insight into the uncertain nature of the very young stellar populations that dominate both these star-forming regions and many galaxies at $z > 6$, which we discuss in Section 6.

For comparison with solar metallicity, we assume a solar gas-phase oxygen abundance of $12 + \log_{10}([O/H]_{\odot}) = 8.69$ (Asplund et al. 2009) unless otherwise noted.¹ For distance calculations and related quantities, we adopt a flat cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. All equivalent widths are measured in the rest-frame and we choose to associate emission with positive values ($W_0 > 0$).

2 SAMPLE SELECTION

While rare in absolute terms, star-forming regions dominated by massive stars at metallicities and ages approaching the conditions found at very high-redshift can be found in appreciable numbers in modern large surveys. Relatively shallow optical spectroscopic observations such as those provided by the Sloan Digital Sky Survey (SDSS) can detect emission from highly ionized gas excited by hard stellar radiation or entrained in dense stellar winds for sufficiently

¹In the context of the Charlot & Bruzual (in preparation) stellar population synthesis models, this is very near the value of the gas-phase oxygen abundance $12 + \log_{10}(O/H) = 8.71$ for $Z = Z_{\odot} = 0.01524$ and near-solar dust-to-metal mass ratio $\xi_d = 0.3$ (see table 2 in Gutkin, Charlot & Bruzual 2016).

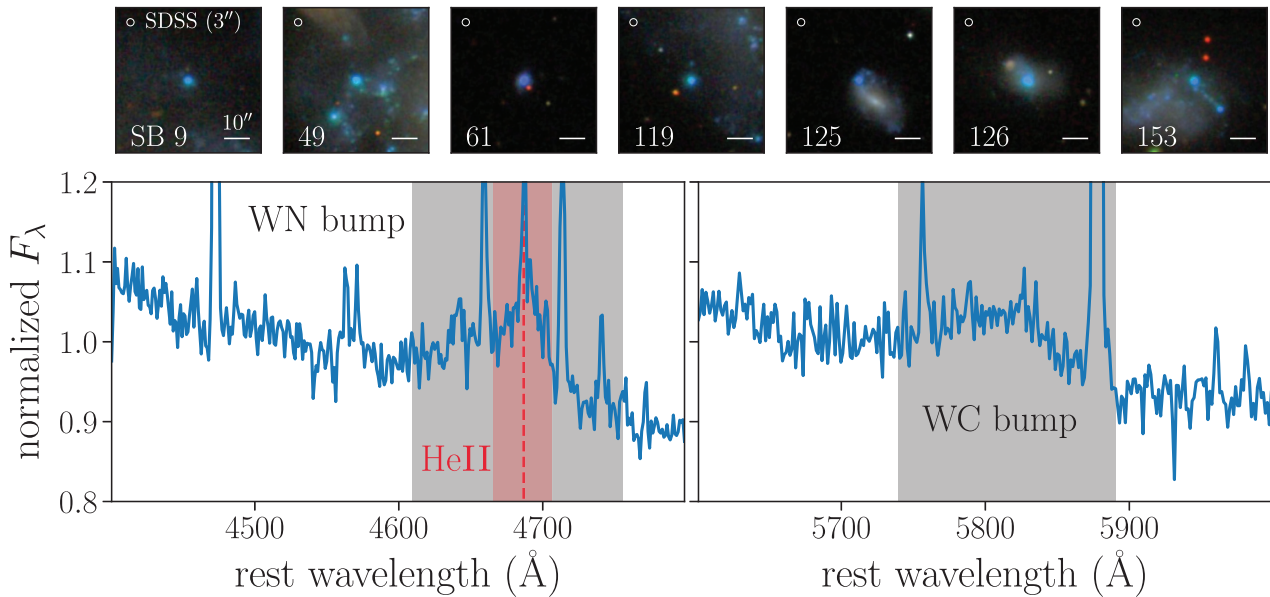


Figure 1. Optical g , r , i cutouts from SDSS centred on our targets, ordered by the identification number from the Shirazi & Brinchmann (2012) sample (top row). A 10 arcsec scalebar and a 3 arcsec diameter circle representing the size of the SDSS spectroscopic aperture are displayed for reference in white over each cutout. On the bottom row, we highlight the broad blue (dominated by WN-type stars with prominent emission from highly ionized nitrogen and helium) and red WR bumps (produced by WC-type spectra with strong ionized carbon lines) in the composite SDSS spectrum of these systems after normalization and median-stacking. The presence of this emission from highly ionized gas entrained in dense stellar winds (in particular, broad He II $\lambda 4686$ clearly visible under the narrow nebular component) is clear evidence that these regions are dominated by massive stars formed in a recent burst.

bright nearby systems. In particular, as recombination lines of He⁺ with an ionization potential of 54.4 eV, the He II $\lambda 4686$ Å optical line and its sibling at 1640 Å in the UV are a powerful signature of galaxies dominated by very hot stars, and broad emission near this line (~ 4600 – 4700 Å) and in the red (~ 5650 – 5800 Å) probe the highly ionized winds of WR stars (e.g. Crowther 2007). These features can be leveraged to efficiently locate the optimum laboratories for the study of prominent populations of young, massive stars.

The Sloan Digital Sky Survey (SDSS) spectroscopic data base has been mined for star-forming galaxies displaying both He II $\lambda 4686$ (Shirazi & Brinchmann 2012, hereafter SB12) and broad WR emission (Brinchmann et al. 2008b, hereafter B08) separately, yielding hundreds of examples of both including some overlap. High-resolution ultraviolet and optical spectra of ten He II emitters from SB12 spanning $7.7 < 12 + \log O/H < 8.5$ were presented in S17. Among the higher metallicity subset at $Z/Z_{\odot} \simeq 0.2$ – 0.5 , these spectra revealed clear stellar wind diagnostics, providing direct insight into the massive stars in these systems. The He II profiles of these systems were dominated by broad emission (FWHM ~ 1600 km s^{−1}) characteristic of populations of envelope-stripped WR and helium stars and very massive stars near the Eddington limit which drive strong, optically thick stellar winds. The youngest of these stellar He II emitters were also found to power nebular C III] at equivalent widths well exceeding that of other systems with less prominent wind emission at these metallicities, suggestive of further potential commonalities with C III] emitters in the reionization era.

The clear detection of optical and UV massive stellar wind signatures alongside nebular gas emission provides a unique opportunity to test stellar population synthesis predictions at the very young ages relevant for modelling galaxies dominated by bursts at high redshift. To establish a statistical sample of UV spectra for galaxies with substantial massive stellar populations to conduct such a test, we

selected an additional seven systems from the SB12 He II emitting sample which also showed prominent WR stellar wind signatures in the optical. In particular, SB12 identified 116 systems with both He II emission and evidence for broad emission at the blue or red WR bumps. These two broad ‘bumps’, located at 4600–4700 and 5650–5800 Å, are composed of blended stellar emission lines of He II, N III, N V, and C III,IV formed in the highly ionized winds of WR stars (e.g. Crowther 2007; Brinchmann et al. 2008b). While the signal-to-noise and resolution of the SDSS spectra render detailed deconvolution of these broad features challenging (though not intractable; e.g. Brinchmann et al. 2004), they can still be clearly detected and distinguished from the narrow nebular emission lines of e.g. [Ar IV], [Fe III], He I, and [N II] interspersed in these regions (Fig. 1). Specifically, we selected galaxies from the star-forming SB12 sample with:

- (i) Clearly detected WR features in the SDSS spectra: specifically, satisfying the SB12 detection quality criterion $WR_{c1} \geq 2$.
- (ii) Sufficiently bright magnitude in the blue for study in a single orbit per grating with *HST*/COS: $u < 18.5$ (aperture magnitude in the 3 arcsec diameter SDSS fibre).
- (iii) Very high specific star formation rate (sSFR), as selected by [O III] $\lambda 5007$ rest-frame equivalent width in-excess of 600 Å.

A total of 11 galaxies satisfied these requirements, three of which were included in the S17 UV sample (SB 80, 179, and 191). We selected the seven highest equivalent width systems of the remaining eight for follow-up with *HST*/COS to complete a UV survey of ten galaxies with extreme WR emission. The new targets are summarized in Table 1, and a summary plot showing SDSS cutouts and their stacked SDSS spectra are displayed in Fig. 1. We will refer to these systems by their identification number from SB12 (prefaced SB) throughout this paper. Note that by construction, our targets are also all present in the earlier sample of SDSS spectra with prominent WR

Table 1. Basic properties of the target star-forming regions, including RA and Dec; redshift; distance and comoving size of the COS aperture (from redshift, corrected for the Local Flow); aperture magnitudes in the SDSS u and i bands; and exposure time in each of the two *HST*/COS gratings. In addition to the seven systems for which new spectra are presented in this paper, we also include the other three systems with identical selection (SB12 He II emitters with optical WR emission) that were analysed in S17 and which we incorporate in our analysis presented here.

SBID	RA (J2000)	Dec (J2000)	z	Distance (Mpc)	2.5 arcsec comoving kpc	u, i AB mag	NUV/G185M, FUV/G160M exposure/s (data set ID)
9	13:04:32.27	-3:33:22.1	0.0045	20	0.24	18.1, 18.7	2260 (LDI204010), 2597 (LDI204020)
49	1:15:33.82	-0:51:31.2	0.0056	31	0.38	17.1, 17.6	2160 (LDI201010), 2597 (LDI201020)
61	14:48:05.38	-1:10:57.7	0.0274	117	1.41	17.9, 17.6	2260 (LDI209010), 2597 (LDI209020)
119	14:32:48.36	9:52:57.1	0.0047	22	0.27	18.1, 18.5	2260 (LDI206010), 2601 (LDI206020)
125	11:32:35.35	14:11:29.8	0.0176	78	0.94	18.2, 18.3	2264 (LDI202010), 2604 (LDI202020)
126	11:50:02.73	15:01:23.5	0.0025	18	0.21	16.7, 17.1	2200 (LDI203010), 2120 (LDI208010)
153	13:14:47.37	34:52:59.8	0.0029	11	0.13	17.6, 18.1	2084 (LDI205010), 2497 (LDI205020)
From S17:							
80	9:42:56.74	9:28:16.2	0.0108	46	0.56	17.9, 18.1	2132 (LCY401010), 2608 (LCY401020)
179	11:29:14.15	20:34:52.0	0.0047	25	0.30	18.3, 18.5	2160 (LCY402010), 2588 (LCY402020)
191	12:15:18.60	20:38:26.7	0.0028	10	0.12	17.7, 18.2	2156 (LCY404010), 2616 (LCY404020)

Table 2. Summary of supplementary MMT observations used to measure [O II] $\lambda 3727$.

Name	Airmass	Integration time (s)
2019 April 9		
SB 125	1.2	3600
SB 126	1.1	3600
SB 119	1.2	1800
SB 153	1.3	1800
2019 October 8		
SB 49	1.1	900

features assembled in B08. We present cross-matched identification numbers from that sample and discuss the targets and their host galaxies in the context of other studies in Appendix A. Together with the targets from S17, our sample comprises the highest sSFR systems with evidence for prominent stellar He II emission in the SDSS spectroscopic survey, ideal for confronting models for very young stellar populations applied at high redshift.

3 OBSERVATIONS AND ANALYSIS

3.1 Optical spectra and imaging

By virtue of their selection from the SDSS Legacy spectroscopic survey (York et al. 2000; Abazajian et al. 2009), high-quality imaging and optical spectra are already available for all seven new targets. Their broad-band photometric magnitudes and the strength of the high equivalent width nebular emission lines which contaminate these bands provide important constraints on the total stellar mass and SFR enclosed in the spectroscopic apertures. We measure aperture photometry centred on each target region from SDSS sky-subtracted and calibrated *ugriz* frames covering each target, adopting a 3 arcsec-diameter aperture corresponding to the size of the SDSS spectroscopic aperture.

The public SDSS spectra for these systems provides $R \simeq 1500$ –2500 (~ 120 –200 km s $^{-1}$) coverage over 3800–9200 Å. As our targets reside at very low-redshifts $z < 0.03$, the [O II] $\lambda\lambda 3727, 3729$ doublet (hereafter [O II] $\lambda 3727$ as we observe it as an unresolved single line) remains outside the SDSS spectroscopic range for all but the

highest redshift, SB 61.² The 3727 doublet is an important component of gas-phase oxygen abundance and ionization parameter determination. Therefore, in addition to the SDSS spectra, we obtained supplementary blue spectra with the Blue Channel spectrograph mounted on the 6.5 m MMT telescope on Mount Hopkins (Table 2). Spectra for four of the targets (SB 119, 125, 126, and 153) were obtained on the night of 2019 April 9 with the 800 gpm grating and 1.5×180 arcsec 2 slit, yielding spectra spanning 3400–5200 Å at lower resolution than SDSS ($R \sim 1200$, corresponding to ~ 250 km s $^{-1}$). Seeing as-measured from the wavefront sensor ranged from 1.1 to 1.5 arcsec over the course of the night. In addition, SB 49 was observed on the night of 2019 October 8 with the 300 gpm grating and 1.0×180 arcsec 2 slit, resulting in coverage of 3500–8500 Å. All observations were reduced using standard longslit data reduction techniques, including wavelength calibration using HeAr/Ne lamp exposures and flux calibration with standard star observations conducted at similar airmass (Feige 34 and BD+33 2642 on April 9, and HZ 14 on October 8). To correct for potential differences in aperture and flux calibration between the SDSS and MMT spectra, we correct the flux of [O II] $\lambda 3727$ measured in the MMT data by the median ratio between other strong lines measured in both spectra. We find a median scatter in this ratio measured between different lines of 8 per cent, which we add in-quadrature to the corrected flux uncertainty.

Nebular line flux and equivalent width measurements were performed using the technique described in more detail in S17. In brief, a base model consisting of a Gaussian emission line (or two in the case of doublets or fitting for multiple components to a single line, as specified) and a linear continuum are fitted to the spectroscopic data in the region of each emission line of interest using the Markov chain Monte Carlo sampling framework EMCEE (Foreman-Mackey et al. 2013). The fits are examined by-eye to ensure that the fit wavelength range is sufficient to accurately capture the continuum level and ensure that the correct features are identified. The resulting posterior distributions after removing a nominal burn-in period straightforwardly provide confidence intervals for the parameters of interest, including fluxes, equivalent widths, and line widths.

²While [O II] $\lambda 3727$ is indeed barely shifted into the wavelength range of the SDSS spectrum for the second highest redshift system SB 125 at $z = 0.018$, the bluest part of the line profile is cut-off at the edge of the SDSS spectrum.

Table 3. Results of fits to the broad-band photometric *ugriz* SED and H β equivalent widths with BEAGLE. These results indicate that the optical photometry of the target systems is in all cases consistent with very young stellar populations formed within the last 3–20 Myr.

Target	$\log_{10}(M/M_{\odot})$	$\log_{10}(\text{SFR}/(M_{\odot} \text{ yr}^{-1}))$	Age/Myr
SB 9	5.05 ± 0.22	-1.95 ± 0.12	$5.5^{+10.8}_{-0.3}$
SB 49	5.80 ± 0.10	-1.20 ± 0.10	$5.5^{+3.5}_{-0.3}$
SB 61	6.95 ± 0.05	-0.05 ± 0.05	$6.2^{+1.4}_{-0.9}$
SB 119	5.12 ± 0.11	-1.88 ± 0.10	$5.6^{+4.0}_{-0.5}$
SB 125	6.66 ± 0.16	-0.65 ± 0.05	$20.2^{+11.4}_{-5.7}$
SB 126	5.81 ± 0.16	-1.42 ± 0.03	$17.1^{+9.2}_{-4.2}$
SB 153	4.93 ± 0.15	-2.07 ± 0.15	$3.7^{+0.3}_{-0.8}$

With this photometry and nebular line information in-hand, important constraints can be placed on the star formation history and ionized gas properties of these star forming regions. As a first step towards constraining the time-scale of recent star formation in these systems, we fit the broad-band photometric measurements and H β equivalent widths measured from the SDSS data. In particular, we use the Bayesian Analysis of GALaxy sEds (BEAGLE, v0.23.0; Chevillard & Charlot 2016) tool to model each system. The underlying stellar models and population synthesis methodology will be discussed in more detail in Section 5 and have been previously described by Gutkin et al. (2016). We assume a constant star formation history (CSFH) with an SMC extinction curve (Gordon et al. 2003) and allow the age of this star formation period ($0 < \log_{10}[t/\text{yr}] < 9$), metallicity ($-2.2 < \log_{10}[Z/Z_{\odot}] < 0$), stellar mass ($0 < \log_{10}[M/M_{\odot}] < 10$), nebular ionization parameter ($-4 < \log U < -1$) and V-band attenuation ($0 < \hat{\tau}_V < 2$) to vary over the indicated uniform prior ranges. We inflate the photometric flux uncertainties by a conservative 5 per cent added in-quadrature to both more heavily weight the spectroscopic H β equivalent width and account for possible differences in spectroscopic versus photometric aperture, and find good agreement with all fit fluxes in the posterior PDFs.

The results of these SED fits are displayed in Table 3. By selection, these systems display extremely high equivalent width optical nebular line emission. Our measurements of their SDSS H β equivalent widths range from 125 to 285 Å, yielding correspondingly strong broad-band nebular line contamination. In particular, the target g-band photometric measurements (contaminated by H β and [O III] $\lambda\lambda 4959, 5007$) are brighter than measurements in the adjacent u-band (close to the continuum, contaminated only by weaker [O II] $\lambda 3727$) by 0.3–1.0 (median 0.5) magnitudes. The results of our SED fitting reflect this, indicating that the photometry of all seven objects is consistent with entirely young stellar populations. The median inferred age of our assumed constant star formation history model range from a very young 3.7 Myr in the most extreme case (SB 153) to ~ 20 Myr for systems with lower equivalent-width optical emission (SB 125, 126), consistent with the presence of WR wind signatures in the optical spectra (Fig. 1). These ages represent an upper limit for the young stellar component of the target systems, as adopting a single-age simple stellar population would require ages uniformly < 10 Myr to reproduce the observed optical line equivalent widths. The stellar masses inferred from this fitting are uniformly low, spanning $10^{4.9} - 10^{7.0} M_{\odot}$, as expected for very young stellar populations with low mass-to-light ratios. This fitting strongly supports the idea that the spectroscopic apertures centred on these systems are entirely dominated by continuum and nebular gas emission from massive stars formed in a very recent star formation episode.

Measurements of the prominent collisionally excited nebular lines provide direct constraints on the ionization state and chemical abundances of this ionized gas. As in S17, we use the software PYNEB (Luridiana, Morisset & Shaw 2015) to derive gas-phase oxygen abundances. We use the temperature-sensitive ratios of [O I] $\lambda\lambda 7320, 7330 / \lambda 3727$ (note the $\lambda\lambda 3726, 3729$ doublet is unresolved by our MMT observations) and [O III] $\lambda 4363 / \lambda 5007$ alongside the density-sensitive [S II] $\lambda 6731 / \lambda 6716$ ratio to derive electron temperatures and densities appropriate for O $^+$ and O $^{2+}$, respectively.³ In the one case for which we lack a measurement of [O II] $\lambda 3727$ (SB 9), we instead estimate $T_e(\text{O II})$ from that measured for O III using the empirical relationship derived by Izotov et al. (2006, equation 14, for the $12 + \log \text{O/H} > 8$ subsample). Attenuation corrections are derived from the Balmer decrement assuming an SMC internal extinction curve (Gordon et al. 2003) and an intrinsic value of H α /H β computed with PYNEB using our derived T_e, n_e after correcting first for Galactic extinction towards each object using the Schlafly & Finkbeiner (2011) maps and the $R_V = 3.1$ extinction curve of Fitzpatrick (1999). As in S17, we do not explicitly correct the Balmer line measurements for underlying stellar absorption as this likely has a negligible effect on the line ratios in the very high nebular equivalent width regime probed by our sample (the ratio of H α /H β is affected at the $\lesssim 2$ per cent level for the ≤ 3 Å corrections typically derived for BCDs; e.g. Izotov, Thuan & Lipovetsky 1997). Additionally, note that any such absorption contamination is self-consistently accounted for in the model comparison with BEAGLE throughout as it is incorporated into the synthetic spectrum from which the model line fluxes are measured.

The resulting metallicities and other optical nebular line properties are presented in Table 4. In addition to high equivalent-width emission indicative of an extremely young effective age, these systems exhibit signs of very highly ionized gas, with dust-corrected $\text{O}_{32} = [\text{O III}] \lambda 4959 + \lambda 5007 / [\text{O II}] \lambda 3727$ values of 5.3–10.7. Their gas-phase metallicities span the regime $8.0 < 12 + \log_{10} \text{O/H} < 8.3$, which assuming a solar abundance of $12 + \log (\text{O/H})_{\odot} = 8.69$ corresponds to a relative metallicity range of $0.2 < Z/Z_{\odot} < 0.4$.

3.2 Ultraviolet spectra

We secured moderate-resolution UV spectra with *HST*/COS for all seven target regions in a program approved in *HST* Cycle 25 (GO:15185, PI:Stark). Our observations were conducted in the NUV with the G185M grating and in the FUV with G160M, with a 2-orbit visit devoted to each object. Each target was first centred with a 7–81 s ACQ/IMAGE exposure with either Mirror A or B, with the optical element and exposure time chosen with the COS ETC based upon an assumed compact source with flat SED in F_{ν} , normalized to the SDSS u-band fibre magnitude. The resulting target acquisition images are shown in Fig. 2, and reveal a range of morphologies at 10–100 pc scales (though note that the marked Mirror B images are complicated by the fainter, displaced secondary image this optical element generates). The remainder of the first orbit of each visit was devoted to G185M observation, with the second orbit occupied by

³We assume atomic transition probabilities for O II from Wiese, Fuhr & Deters (1996), Froese Fischer & Tachiev (2004) and effective collision strengths from Tayal (2007). For O III we adopt the atomic data of Storey & Zeippen (2000), Froese Fischer & Tachiev (2004) and collision strengths of Storey, Sochi & Badnell (2014). And for S II, we utilize the transition probabilities presented by Podobedova, Kelleher & Wiese (2009), Tayal & Zatsarinny (2010) and collision strengths from Tayal & Zatsarinny (2010).

Table 4. Optical nebular line properties measured from SDSS and MMT spectra where available. Note that since we do not have a measurement of [O II] $\lambda 3727$ for SB 9, the metallicity for this system (\dagger) was determined using only [O II] $\lambda 7320$, 7330.

Name	O ₃₂	R ₂₃	[O III] $\lambda 5007$ W ₀ (Å)	H β W ₀ (Å)	E(B–V)	T _e (O III) 10 ⁴ K	12 + log ₁₀ (O/H)
SB 9	–	–	920 ± 52	209 ± 9	0.04	1.04 ± 0.02	8.30 ± 0.05 [†]
SB 49	6.9 ± 0.9	6.3 ± 0.2	1001 ± 39	223 ± 5	0.21	0.98 ± 0.02	8.20 ± 0.04
SB 61	10.7 ± 0.6	10.4 ± 0.5	1127 ± 82	146 ± 8	0.11	1.29 ± 0.02	8.11 ± 0.04
SB 119	5.5 ± 0.5	7.4 ± 0.4	947 ± 54	208 ± 15	0.13	1.10 ± 0.02	8.11 ± 0.04
SB 125	8.6 ± 0.7	9.2 ± 0.3	824 ± 34	125 ± 4	0.25	1.13 ± 0.02	8.19 ± 0.03
SB 126	8.6 ± 0.6	7.7 ± 0.3	723 ± 36	134 ± 4	0.18	1.21 ± 0.02	8.02 ± 0.04
SB 153	5.3 ± 0.7	9.1 ± 0.4	1719 ± 133	286 ± 12	0.00	1.12 ± 0.02	8.20 ± 0.04
From S17:							
SB 80	3.7 ^{+0.3} _{-0.3}	9.6 ^{+1.2} _{-1.1}	1195 ± 132	243 ± 17	0.13	1.15 ± 0.02	8.24 ± 0.06
SB 179	2.7 ^{+0.1} _{-0.1}	8.5 ^{+1.9} _{-1.4}	770 ± 81	196 ± 8	0.17	1.07 ± 0.02	8.35 ± 0.07
SB 191	9.5 ^{+0.7} _{-0.9}	9.0 ^{+1.8} _{-1.5}	1649 ± 239	393 ± 23	0.02	1.05 ± 0.03	8.30 ± 0.07

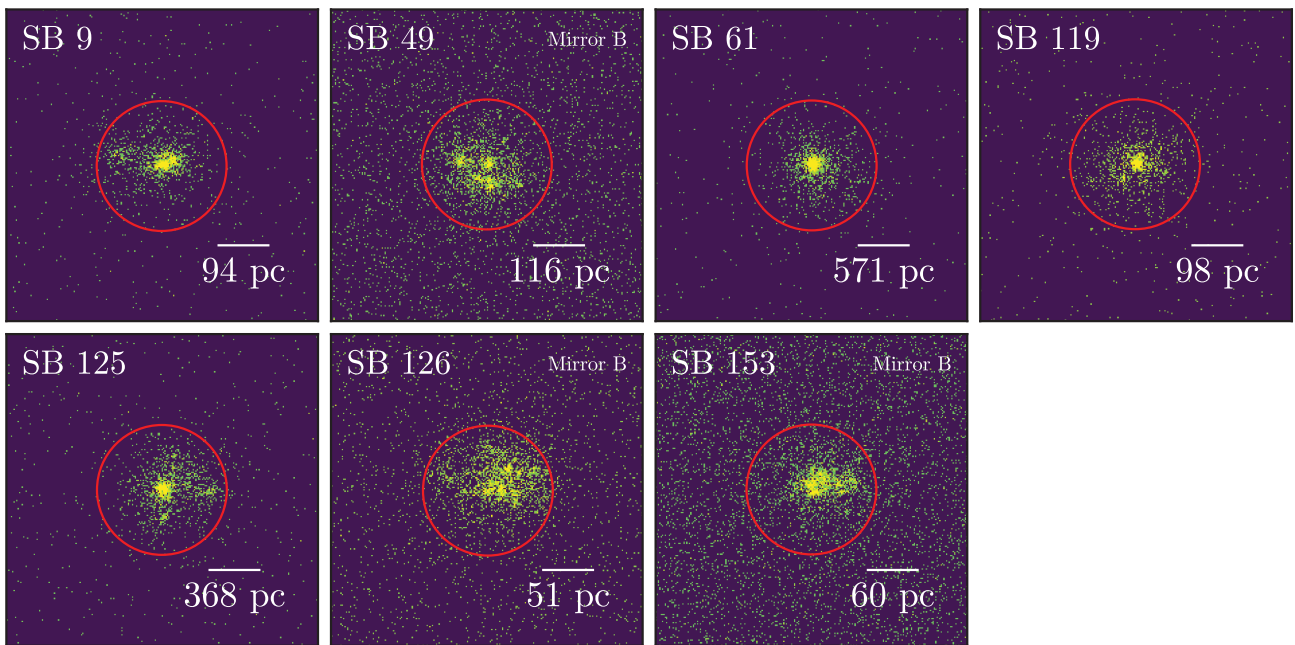


Figure 2. Target acquisition images taken in the NUV with *HST/COS*. A 1.0 arcsec scalebar is drawn in each case, labelled with the comoving distance this angle corresponds to at the estimated distance of each object. Note that the brightest three sources were observed with Mirror B to ensure instrument safety, which produces a fainter second image at about half-intensity displaced by 20 pixels from the primary image. Thus, some of the structure visible in these images (noted) is artificial. While each object is clearly dominated by a bright central object, the additional diffuse flux and fainter compact sources visible around each suggest that we are viewing light from a complex of star-forming regions.

G160M. The central wavelength of each observation was selected with the SDSS redshift in mind to ensure that the full stellar C IV $\lambda\lambda 1548, 1550$ profile, He II $\lambda 1640$, and O III] $\lambda\lambda 1661, 1666$ were located in G160M segments A and B, and that the [C III] $\lambda 1907 +$ C III] $\lambda 1909$ doublet (hereafter C III] $\lambda 1908$) was centred in the middle segment NUVB.

Two visits of our program required repeat observation. On 2018 April 29, guide star tracking was lost during reacquisition at the beginning of the second orbit for SB 126. While the G185M integration was successfully completed before fine guidance loss occurred, this prevented the G160M integration (LDI203020) from proceeding. A single-orbit repeat was scheduled as-per HOPR 90495, and a G160M spectrum was successfully recorded for this object on 2018 May 27 (LDI208010). During the original scheduled visit targeting

SB 61 on 2018 June 20, the FGS acquired the guide stars after target acquisition had been attempted. As a result, the aperture was not correctly centred on-target for the spectroscopic exposures, resulting in an unexpectedly low continuum level and signal-to-noise in both the G160M and G185M spectra (LDI207020 and LDI207010, respectively). These observations were repeated following HOPR 90852 with a successful target acquisition on 2018 August 9, yielding spectra that reach the expected S/N which we use in this paper (LDI209020 and LDI209010).

All spectroscopic observations were taken in time-tagged photon-address mode (TIME-TAG) with wavelength calibration lamp exposures (FLASH=YES) to allow for optimal data correction post-observation. We cycle through all four focal plane offset positions for each grating (FP-POS=ALL) to mitigate the effect of fixed pattern

Table 5. Line measurements from the new ultraviolet *HST/COS* spectra presented in this paper along with those for the three other extreme WR galaxies included in [S17](#). In particular, we present equivalent widths integrated over two broad regions capturing the C IV resonant stellar P-Cygni absorption trough and broad He II stellar wind emission; and flux and equivalent width measurements for nebular O III] and C III] emission. Where these lines are undetected, we present 2.5σ upper limits thereon. Nebular lines marked by a † were contaminated by MW Al II $\lambda 1670.79$ absorption and either entirely suppressed or likely significantly affected, and should be interpreted with caution. We also present C/O estimates from population synthesis fits to the UV and optical nebular lines as described in Section 5.1 and adopted in the remainder of the fits presented in this paper.

Name	C IV $\lambda 1549$ P-Cygni Absorption, $W_0 \text{ \AA}^{-1}$	He II $\lambda 1640$ Emission, $W_0 \text{ \AA}^{-1}$	O III] 1661 $10^{-15} \text{ ergs s}^{-1} \text{ cm}^{-2}$ ($W_0 \text{ \AA}^{-1}$)	O III] 1666 $10^{-15} \text{ ergs s}^{-1} \text{ cm}^{-2}$ ($W_0 \text{ \AA}^{-1}$)	C III] 1908 $10^{-15} \text{ ergs s}^{-1} \text{ cm}^{-2}$ ($W_0 \text{ \AA}^{-1}$)	$\log_{10}(\text{C/O})$ (Section 5.1)
SB 9	-4.40 ± 0.07	2.16 ± 0.16	0.59 ± 0.06 (0.32 ± 0.03)	1.89 ± 0.11 (1.05 ± 0.07)	6.61 ± 0.44 (4.86 ± 0.35)	-0.36 ± 0.02
SB 49	-6.51 ± 0.04	3.20 ± 0.10	<0.97 (<0.28) †	4.50 ± 0.23 (1.46 ± 0.08)	6.86 ± 0.41 (3.12 ± 0.20)	-0.67 ± 0.06
SB 61	-2.28 ± 0.09	1.99 ± 0.18	1.19 ± 0.09 (0.60 ± 0.05)	4.12 ± 0.11 (2.19 ± 0.07)	15.16 ± 0.80 (9.43 ± 0.52)	-0.29 ± 0.02
SB 119	-4.26 ± 0.10	4.27 ± 0.24	0.71 ± 0.05 (0.57 ± 0.04)	0.79 ± 0.04 (0.65 ± 0.03)	4.24 ± 0.47 (3.79 ± 0.45)	-0.36 ± 0.03
SB 125	-4.62 ± 0.09	2.67 ± 0.21	<0.47 (<0.30)	1.50 ± 0.09 (1.05 ± 0.07)	6.11 ± 0.39 (4.80 ± 0.33)	-0.29 ± 0.04
SB 126	-3.96 ± 0.03	1.93 ± 0.07	<1.24 (<0.19)	<1.24 (<0.20) †	11.68 ± 0.62 (2.44 ± 0.13)	-0.29 ± 0.06
SB 153	-3.75 ± 0.06	3.94 ± 0.13	1.11 ± 0.09 (0.43 ± 0.04)	1.20 ± 0.06 (0.50 ± 0.03) †	17.69 ± 0.41 (8.17 ± 0.23)	-0.30 ± 0.03
From S17 :						
SB 80	-3.12 ± 0.13	3.29 ± 0.27	0.60 ± 0.04 (0.55 ± 0.04)	1.69 ± 0.07 (1.68 ± 0.08)	<4.1 (<4.0)	-0.48 ± 0.27
SB 179	-5.22 ± 0.11	4.67 ± 0.23	<1.1 (<0.9)	1.36 ± 0.06 (1.17 ± 0.06)	6.80 ± 0.23 (8.71 ± 0.42)	-0.24 ± 0.04
SB 191	-5.20 ± 0.08	4.02 ± 0.16	1.19 ± 0.07 (0.64 ± 0.04)	$-$ ($-$) †	15.91 ± 0.39 (11.33 ± 0.34)	-0.31 ± 0.01

noise. The data were reduced with the default settings for calibration and extraction using CALCOS v3.3.5 and CRDS v7.3.0 (*HSTDP* v2019.2). The final 1D spectra, with resolution elements (resels) of 73.4 m\AA in the FUV/G160M and 102 m\AA in the NUV/G185M, achieve a typical effective spectral resolution of 0.25 \AA as-inferred from the width of Milky Way absorption lines (identical to that found in [S17](#)) and median signal-to-noise of 11 (4) per resel at 1450 \AA (1700 \AA) in G160M and 1.7 per resel at 1900 \AA in G185M.

We measure emission lines in the *HST/COS* spectra in the same manner as for the optical emission lines. The high spectral resolution afforded by the G160M and G185M gratings allows us to easily identify and separate Milky Way (MW) and ISM absorption lines, alleviating blending concerns except in the most extreme cases. In addition to our measurement of O III] and C III] semiforbidden nebular line emission, we define two integration regions designed to measure the strength of the two key stellar wind features accessed by the G160M spectra. In order to ensure that ISM and MW absorption lines do not contaminate these measurements, we first mask-out the spectra $\pm 1.5 \text{ \AA}$ around all identified strong absorption lines and interpolate the spectrum over these windows. We estimate the equivalent width of the C IV P-Cygni wind absorption trough by integrating this cleaned spectrum between 1530 and 1550 \AA relative to a linear continuum determined by evaluating the median flux in two windows chosen to avoid wind and nebular line features: $[1500:1525]$ and $[1565:1575]$. Likewise, we measure the strength of emission in the He II $\lambda 1640$ wind and nebular emission line in the window 1630 – 1650 \AA relative to a continuum estimated from flux at $[1620:1630]$ and $[1650:1660]$.⁴ We present the resulting measurements in Table 5.

The measured properties of the UV spectra of our sample (Fig. 3) further evince a dominant population of massive stars in these galaxies. We detect C III] emission in all seven systems, at equivalent widths ranging from 2 – 9 \AA and O III] $\lambda 1666$ in all but one (SB 126, where it is fully absorbed by the MW Al II $\lambda 1671$ line). Including the three additional [S17](#) objects reveals a similar detection rate, extending the C III] equivalent width range up to just over 11 \AA (in SB 191) with only one system undetected in this doublet (SB 80).

⁴We use only the blueward continuum window to integrated He II $\lambda 1640$ for SB 80 from [S17](#) as strong Al II $\lambda 1670$ absorption contaminates a significant portion of the redward window.

The essentially unanimous detection of this doubly ionized carbon and oxygen emission in the UV confirms the presence of a strong ionizing continuum at ~ 20 – 40 eV .

As mentioned, the other high-ionization line complexes in our G160M spectra provide a direct window on to the massive stars powering this nebular gas emission. The C IV $\lambda\lambda 1548, 1550$ doublet complexes show no clear sign of narrow nebular emission, instead revealing a characteristic P-Cygni profile consisting of broad redshifted emission and blueshifted absorption extending down to ~ 1530 – 1535 \AA (~ 2500 – 3500 km s^{-1}), reaching equivalent widths in absorption of 2.3 – 6.5 \AA . This feature is formed in the dense winds of massive O stars, and is observed to extend to comparable terminal velocities in LMC stars at $Z/Z_{\odot} \simeq 0.5$ (e.g. Massey et al. 2004; Crowther et al. 2016). Likewise, rather than the nebular gas emission typically observed in He II $\lambda 1640$ at metallicities below $\sim 0.2Z_{\odot}$, the target systems all show broad emission with FWHM ~ 1500 – 2000 km s^{-1} . This emission is characteristic of massive WN and O If stars, and is formed in extremely dense and optically thick stellar winds driven by stars approaching the Eddington limit. The prominence of this emission is particularly extraordinary, with three systems exceeding 4 \AA . We will discuss the context and implications of these detections in more detail in the following section.

4 THE ULTRAVIOLET STELLAR WIND FEATURES

Imprints of extremely hot stellar atmospheres dominate the UV spectra of these 10^5 – $10^7 M_{\odot}$ star-forming complexes. These features represent a powerful probe of the massive stars populating these galaxies and an opportunity to test and inform stellar population synthesis models at very young effective ages. Before proceeding to describe the experiment, we will conduct leveraging these stellar wind profiles in more detail, in this section we will first explore their strength in our sample relative to other star-forming galaxies and determine whether modern population synthesis models approach their equivalent width distribution at all.

While obtaining profiles at this signal-to-noise for individual systems at cosmological redshifts is extremely challenging, prominent stellar wind signatures are routinely encountered in the integrated light spectra of galaxies dominated by recent star formation. Indeed, C IV P-Cygni and broad He II emission have been detected in

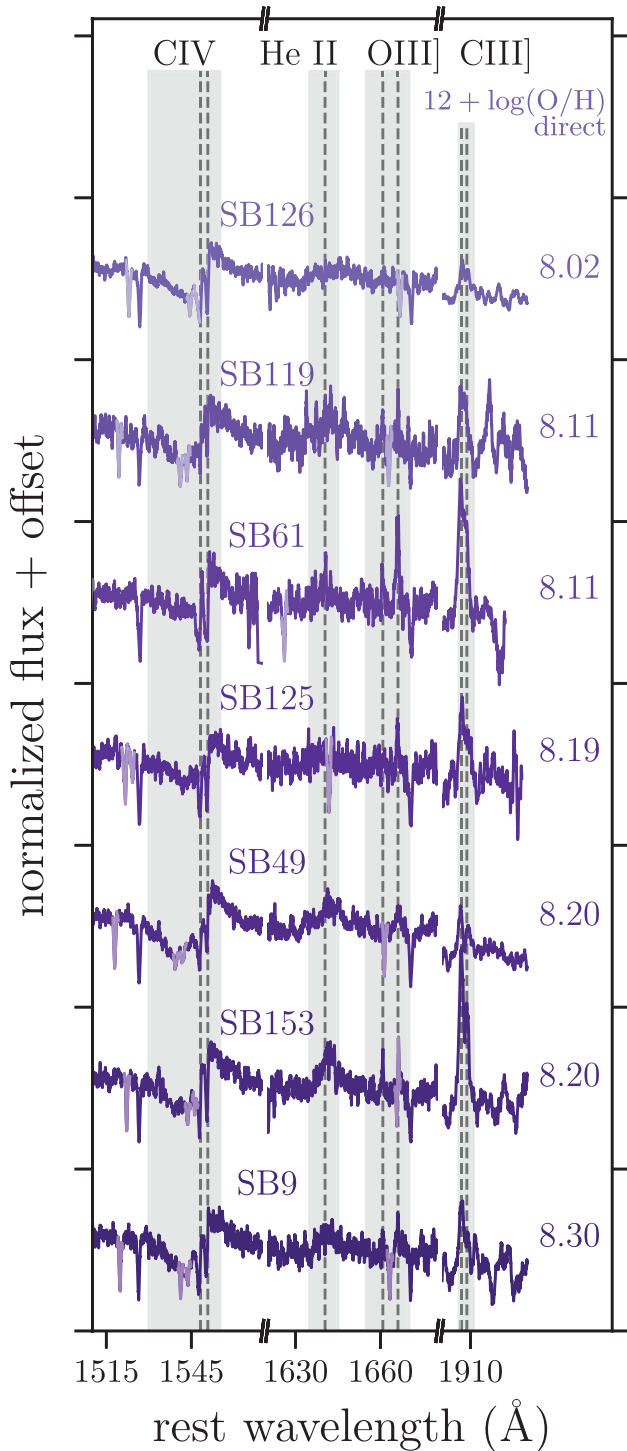


Figure 3. Ultraviolet *HST/COS* spectra for the seven new star-forming regions targeted in this paper. Sections of the spectra impacted by intervening Milky Way absorption lines are plotted in a lighter shade. The G160M+G185M data reveal O III] and C III] nebular emission in addition to very prominent stellar wind signatures both in resonant C IV P-Cygni profiles and broad wind emission in He II.

gravitationally lensed and stacked spectra of galaxies out to $z \sim 4$ (e.g. Shapley et al. 2003; Dessauges-Zavadsky et al. 2010; Erb et al. 2010; Jones, Stark & Ellis 2012; Steidel et al. 2016), at strengths that have challenged the accuracy of canonical stellar population synthesis

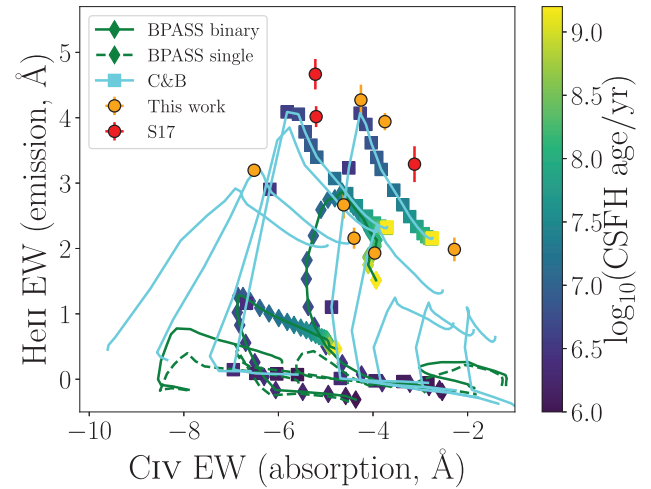


Figure 4. The strength of stellar He II emission versus C IV stellar absorption for our extreme WR galaxies and selected population synthesis models. For each set of models, we plot tracks of age assuming constant star formation for different stellar metallicities. Each track rises to peak at a maximal He II equivalent width at an age corresponding to the dominant source of this optically thick wind emission (canonical massive WR/OIF stars or QHE binary products; ages are displayed in colour-coded points for a subset), and lower metallicity model tracks tend to reside farther to the right in median value (weaker C IV P-Cygni absorption due to weaker stellar winds overall). The observed galaxies (orange and red points) reach broad stellar He II emission equivalent widths well in-excess of canonical model prediction (e.g. the BPASS single-star models, green dashed). However, comparison with BPASS and C&B models suggests that accounting for rapidly rotating stars produced by binary mass transfer or adopting new predictions for the evolution and atmospheres of very massive stars can both boost these predictions close to the equivalent width regime observed.

models. These $z \simeq 2-3$ galaxies reach stellar He II $\lambda 1640$ equivalent widths of $1-3 \text{ \AA}$, often significantly in-excess of canonical stellar population synthesis predictions without invoking solar or supersolar stellar metallicities (e.g. Leitherer et al. 1999; Brinchmann, Pettini & Charlot 2008a).

The star-forming regions assembled here reveal even more prominent stellar wind emission, and are among the most intense stellar He II-emitters known. In Fig. 4, we plot the equivalent width of the stellar He II emission against that of the P-Cygni C IV absorption trough, measured as described in Section 3.2, for the full sample of extreme WR galaxies in Table 5. All of the galaxies in this sample power stellar He II emission at equivalent widths $\geq 2 \text{ \AA}$, reaching 4.7 \AA in the most extreme case (SB 179). The only comparable or more prominent He II emission in integrated light spectra has been found in similar star-forming complexes at low redshift; in particular, NGC 3125-1 (7.1 \AA ; Chandar, Leitherer & Tremonti 2004; Hadfield & Crowther 2006; Wofford et al. 2014), SSC-N in II Zw 40 (7.1 \AA ; Leitherer et al. 2018), NGC 5253-#5 (5.1 \AA ; Smith et al. 2016), and the coadded spectrum of stars in the centre of R136 in the LMC (4.5 \AA ; Crowther et al. 2016). This powerful though not unprecedented emission places our targets uniformly in-excess of the predictions of canonical stellar population synthesis models (e.g. Brinchmann et al. 2008a), which even including bursts struggle to exceed 2 \AA in the 1640 line.

However, recent developments in stellar population modelling dramatically alter this picture. Canonically, broad He II emission in integrated galaxy spectra is produced primarily by very massive OIF and classical WR stars; initially very massive ($>25 M_{\odot}$) stars

which remove their own envelopes through intense stellar wind mass-loss exposing their He-burning core after evolving off the main sequence (e.g. Crowther 2007). In this case, prominent He II emission appears only for a short period ~ 5 Myr following an instantaneous burst of star formation. However, wind emission in He II is a common trait of very hot stellar cores without a thick envelope of hydrogen, and such stars can also be produced in abundance by binary interaction. Both mass donor stars stripped of their outer envelopes and mass gainers rejuvenated and driven to high rotation rates by conservative mass transfer can evolve into helium stars, potentially driving strong wind emission in He II (e.g. Cantiello et al. 2007; Götzberg et al. 2018). In particular, Eldridge & Stanway (2012) demonstrate that including stars spun-up to very high rotation rates by mass accretion which become fully mixed and consequently undergo quasi-homogeneous evolution (QHE; e.g. Yoon & Langer 2005) can substantially impact on the strength of this wind emission. They find that incorporating a prescription for generating such stars into their Binary Population And Spectral Synthesis (BPASS) code brings model predictions assuming a constant star formation history into much better agreement with observations at $z \sim 2-3$.

Changes to the modelling of very massive single stars can also significantly alter predictions for stellar He II and other high-ionization wind lines. In addition to massive stars which evolve through a WR phase during core helium-burning, very massive stars on the MS but so luminous as to reside close to the Eddington limit can also drive dense winds with He II in prominent emission, dominating He II emission in very young (< 5 Myr) star clusters (Olf/WNh stars; e.g. Crowther et al. 2016). The latest version of the Bruzual & Charlot (2003) stellar population synthesis models (Charlot & Bruzual, in-preparation; hereafter C&B) adopt new prescriptions for the evolution, mass-loss rates, and atmospheres of massive stars (as described in Section 5.1 and Vidal-García et al. 2017) which together significantly affect the strength of optically thick wind lines like He II. In particular, the underlying PARSEC stellar evolutionary models for massive stars adopt a new Eddington-factor formalism to predict mass-loss rates for the most massive stars, which allows strong stellar winds to be driven at lower metallicities and very young ages for sufficiently luminous stars (Chen et al. 2015).

It is instructive to compare the predictions of these modern stellar population synthesis codes for the UV stellar wind features in the context of our new measurements. In Fig. 4, we plot both BPASS (v1.1, with and without binary interaction included⁵) and C&B model tracks for comparison with our *HST*/COS stellar wind measurements. For each set of models, we plot tracks in age assuming a constant star formation history at the relevant stellar metallicities provided by the models ($0.05 < Z/Z_{\odot} < 1.0$ for BPASS; and for C&B, $0.16 \leq Z/Z_{\odot} \leq 1$). Each fixed-metallicity track (line) spans a limited range of C IV equivalent widths, with higher metallicity models with stronger O-star winds presenting deeper C IV absorption on average. Each track generally begins at low-EW in He II and first moves nearly horizontally to increasingly strong C IV absorption as the O-star population is built-up. The tracks then progress to higher He II equivalent widths which peak at a certain age before settling-down to a more moderate equilibrium value in He II and C IV as the continuum contribution of somewhat older stellar populations with weaker winds is built-up.

⁵In particular, we plot the results for BPASS v1.1 for consistency with Eldridge & Stanway (2012), but note that we measured He II $\lambda 1640$ in the latest v2.2 results and did not find significantly higher equivalent widths than reported in v1.1.

As can be seen for the selected tracks for which we plot individual model points with colour-coded ages, this peak value in He II occurs at different ages for different model assumptions. In particular, the BPASS models reach this value at $\sim 10-50$ Myr as the contribution from initially lower mass stars $\sim 10-20 M_{\odot}$ that are spun-up into QHE by mass transfer reach a maximum. These models with QHE boost He II substantially over the BPASS single-star predictions, and this effect is most prominent at $Z/Z_{\odot} = 0.2$ due to the competing effects of increasing numbers of stars undergoing QHE at lower metallicities and the diminishing strength of their stellar winds (Eldridge & Stanway 2012). While this does push the models closer to the observed population, intriguingly at a similar range of C IV equivalent widths as observed for our strongest He II emitters, the maximum He II equivalent widths reached by these models still fall short of the most extreme emitters by at least 1 \AA .

In contrast, the C&B models reach a maximum in He II at much younger ages. Because the dominant source of the stellar He II emission in these single star models is initially massive Olf/WNh and He-burning WR stars rather than a broader variety of stars spun-up by binary interaction, the tracks boost to their respective maxima at $\sim 4-5$ Myr rather than > 10 Myr. Additionally, both the peak and equilibrium values of the He II equivalent width for a given track are uniformly greater than predicted by the other models, extending to 3–4 and 2–3 \AA , respectively. As in the BPASS predictions, the maximum in He II wind emission is reached at an intermediate metallicity rather than at solar; in the case of the C&B models, this occurs at $Z = 0.006-0.008$, or $Z/Z_{\odot} \simeq 0.5$ (approximately the metallicity of the LMC). While the tracks still fall slightly short of the most intense emission observed (and more significantly in comparison to the most extreme local systems mentioned above; e.g. Wofford et al. 2014; Leitherer et al. 2018), these models have now entered a similar part of this stellar wind parameter space as that occupied by the star-forming regions presented in this work.

The C&B models come close to spanning the stellar C IV-He II equivalent width regime that our target galaxies occupy. This suggests that the improved treatment of the evolution of single massive stars and their winds has alleviated at least some of the tension previously noted between predicted and observed stellar He II strengths in very young star-forming systems. However, these models do not include binary mass transfer, mergers, or initially high rotation rates, which the BPASS models discussed above demonstrate could substantially boost wind emission in He II as well. The impact of these processes on the evolution of the most massive stars remains highly uncertain, yet increasingly relevant to our ability to interpret high-redshift galaxies dominated by very young stellar populations. In the following section, we describe an experiment testing whether the C&B models can self-consistently reproduce the strong C IV and He II wind lines alongside the nebular emission lines powered by the same massive stars. By searching for discrepancies with the latest single-star population synthesis predictions, we will determine whether the ionizing spectrum and stellar wind lines are impacted significantly by binary evolution or other uncertain physics relevant at these young inferred stellar ages. The results of this experiment will illustrate the impact that ignoring such processes may have on the validity of spectral inference for galaxies undergoing bursts of star formation in the distant Universe.

5 UV–OPTICAL SPECTRAL SYNTHESIS

As outlined in the sections above, our high-quality UV–optical spectroscopy access powerful diagnostics of both the winds driven by massive stars and the ionizing radiation fields they produce by way of

reprocessed nebular emission. The very high equivalent width of the optical lines these systems power implies a dominant young $\lesssim 10$ Myr stellar population, comparable to the extremely high sSFR systems glimpsed in the reionization era (e.g. Roberts-Borsani et al. 2016; Endsley et al. 2021) and probing conditions where the uncertain treatment of very massive stars in population synthesis will have an outsized impact on our ability to correctly interpret and model observations.

As discussed in Section 4, in this paper we will focus on comparisons to the latest predictions of the C&B stellar population synthesis framework. While these state-of-the-art single-star evolution models come far closer to reproducing the observed distribution of stellar wind line strengths than previous generations, they still do not incorporate the impact of mass transfer, mergers, or rapid rotation. Systematic issues in reproducing the spectra of very young star-forming regions such as these can provide valuable insight into the physics at work in shaping the evolution of the most massive stars.

In particular, when the strong nebular emission line fluxes are fitted with the C&B models and used to constrain the properties of the underlying stellar population, we are interested in whether the strength of the UV stellar wind features of the responsible massive stars can be successfully reproduced. If the C&B models are unable to match both sets of features, the discrepancies may reveal the impact of stellar interactions commonly neglected in population synthesis at these metallicities and ages. We describe our methodology and the results of these fits in the following two subsections.

5.1 Methodology

We will conduct this comparison experiment against the updated C&B single-star models using the BEAGLE inference framework. As implied by our findings in Section 4, the improvements in the treatment of very massive star evolution adopted by these models has alleviated some of the tension previously found with the stellar He II wind line; but these models still neglect the physics of binary evolution and rotation, which may have a significant impact even at very young stellar population ages. The version of the C&B models used in this work is described in more detail in Vidal-García et al. (2017, their Section 2.1 and Appendix A, updating Gutkin et al. 2016). The most important features of this code for the purposes of this work concern the revised predictions for the evolution and atmospheres of the most massive stars (extended now to $300 M_{\odot}$). The latest predictions from the Padova and TRIeste Stellar Evolution Code (PARSEC) are now adopted to predict the evolution of individual massive stars (Bressan et al. 2012; Chen et al. 2015). Notably, these latest PARSEC results include a revised formalism for the prediction of stellar wind driven mass-loss rates based upon the stellar Eddington factor (Vink et al. 2011). This is in contrast to most other stellar evolutionary codes which predict O and WR mass-loss rates using older calibrations (Vink, de Koter & Lamers 2001) derived under assumptions about wind driving which likely break down for stars at the highest luminosities and subsolar metallicities (see e.g. Gräfener & Hamann 2008; Müller & Vink 2008; Sander, Vink & Hamann 2020). In addition, the C&B models leverage the large library of non-LTE, line-blanketed, spherically expanding WR model atmospheres produced by the Potsdam Wolf-Rayet group (PoWR) spanning metallicities from $Z = 0.001$ – 0.014 , effective temperatures from $10^{4.4}$ – $10^{5.3}$ K, and a range of CNO mass fractions⁶ (Gräfener, Koesterke & Hamann 2002; Hamann & Gräfener 2003;

Hainich et al. 2014, 2015; Sander et al. 2015; Todt et al. 2015). These updated mass-loss and atmosphere prescriptions for massive stars near the Eddington limit have significantly improved bulk agreement with observations for stellar lines like He II $\lambda 1640$ produced in the optically thick winds such OIF and WR stars drive, as we discussed in Section 4 above.

The spectral inference framework BEAGLE (Chevallard & Charlot 2016, introduced in Section 3.1) is readily able to fit spectrophotometric observables using the C&B stellar population models and associated CLOUDY nebular line predictions, and provides the inference framework we require for this experiment. For this analysis, we choose to adopt the fiducial Chabrier (2003) initial mass function (IMF) with upper mass cutoff of $300 M_{\odot}$ (as we will discuss below, the contribution of very massive stars $> 100 M_{\odot}$ to high-ionization stellar wind lines can be dominant; e.g. Crowther et al. 2016). For all fundamental variables describing the stellar population, we adopt a uniform prior over the adopted range for simplicity. Since the initial photometric fits including nebular emission were well-fit by a CSFH model (Section 3.1) and this likewise populates the region of stellar wind parameter space occupied by our observations (Fig. 4), we parametrize our assumed SFH as such. However, in addition to the variable start time (which we allow to vary from $10^{6.1}$ – $10^{9.0}$ yr) and stellar mass ($M \leq 10^{10} M_{\odot}$) for this recent period of star formation, we allow an early truncation of this constant star formation history up to 10 Myr before the present time to allow additional flexibility in the mix of very massive stars present. In other words, we do not force the fit star formation episode to continue all the way to the present day, and effectively allow it to shrink in duration towards an instantaneous burst within the last 10 Myr. We allow the stellar metallicity to freely vary with $-2.2 \leq \log_{10}(Z/Z_{\odot}) \leq 0$. While this framework explicitly couples the interstellar metallicity to this stellar metallicity, depletion of refractory elements on to dust grains can change the gas-phase oxygen abundance at fixed Z . We allow the dust-to-metal mass ratio ξ_d to vary from 0.1–0.5, which effectively corresponds to a factor of 1.6 range in the ratio of gas-phase oxygen to stellar iron (Gutkin et al. 2016). This ratio plays an important role in modulating the joint nebular and stellar spectra of star-forming systems (e.g. Steidel et al. 2016; Strom et al. 2018; Sanders et al. 2020), since oxygen is the primary species probed in the nebular spectrum while the winds and ionizing spectra of massive stars are shaped primarily by iron. Thus a model preference for low ξ_d can be interpreted in part as a preference for enhanced α/Fe in this context. Since we are studying relatively small star-forming regions, we assume an SMC extinction curve with variable $0 \leq \tau_V \leq 2$. Finally, we allow the nebular ionization parameter to vary over $-4 \leq \log U \leq -1$.

In this experiment, we are interested first in fitting the nebular line fluxes which constrain the stellar ionizing spectrum and gas-phase metallicity. In particular, we focus on the set of strong optical nebular lines necessary to measure the gas-phase oxygen abundance via the direct method: [O II] $\lambda 3727$, [O III] $\lambda 4363$, H β , [O III] $\lambda \lambda 5007$, and H α . Since the gas densities inferred from the [S II] $\lambda \lambda 6716, 6731$ doublet are all consistent with 100 – 200 cm^{-3} , for simplicity we exclude these lines from the fits and instead fix the gas density to the nearest grid value of 100 cm^{-3} . For additional leverage on the dust attenuation, we add H γ ; and to better anchor the recent star formation history of the models, we also fit equivalent widths measured in SDSS for H γ , H β , [O III] $\lambda 5007$, and H α . In each case, the observed line flux or equivalent width is compared to the flux measured via simple integration using a linear continuum in the model spectra. We choose the integration region and continuum windows for each line manually to ensure that they capture the entire line flux and that the continuum is not biased by other features. Since

⁶see <http://www.astro.physik.uni-potsdam.de/~wrh/PoWR/powrgrid1.php>.

Table 6. Constraints on model parameters of interest from our BEAGLE fits to nebular line emission, with or without including the UV stellar equivalent width measurements in the fit. In particular, we highlight the stellar mass, SFR, and beginning and end of the CSFH period; as well as the gas phase oxygen abundance, stellar metallicity, dust-to-metal mass ratio, and ionization parameter.

Target	$\log_{10}(M/M_{\odot})$	$\log_{10}(\text{SFR}/M_{\odot}\text{yr}^{-1})$	t_1	t_0	$12 + \log_{10}(\text{O}/\text{H})$	$\log_{10}(Z_{\star}/Z_{\odot})$	ξ_d	$\log_{10}U$
Optical nebular lines only								
SB 9	$5.29^{+0.05}_{-0.05}$	$-1.95^{+0.06}_{-0.03}$	$7.2^{+0.1}_{-0.1}$	<6.0	$8.39^{+0.04}_{-0.04}$	$-0.37^{+0.07}_{-0.06}$	$0.29^{+0.13}_{-0.13}$	$-2.95^{+0.08}_{-0.06}$
SB 49	$6.26^{+0.04}_{-0.04}$	$-0.93^{+0.04}_{-0.04}$	$7.2^{+0.0}_{-0.0}$	<6.0	$8.53^{+0.03}_{-0.03}$	$-0.30^{+0.04}_{-0.03}$	$0.16^{+0.06}_{-0.04}$	$-2.74^{+0.06}_{-0.06}$
SB 61	$7.56^{+0.06}_{-0.08}$	$-0.05^{+0.05}_{-0.03}$	$7.6^{+0.1}_{-0.1}$	<6.0	$8.25^{+0.04}_{-0.07}$	$-0.57^{+0.05}_{-0.07}$	$0.13^{+0.04}_{-0.02}$	$-2.06^{+0.06}_{-0.07}$
SB 80	$6.10^{+0.04}_{-0.04}$	$-0.91^{+0.04}_{-0.04}$	$6.9^{+0.0}_{-0.1}$	<6.0	$8.33^{+0.03}_{-0.03}$	$-0.47^{+0.08}_{-0.04}$	$0.21^{+0.15}_{-0.08}$	$-2.59^{+0.05}_{-0.05}$
SB 119	$5.60^{+0.06}_{-0.07}$	$-1.71^{+0.07}_{-0.04}$	$7.3^{+0.1}_{-0.1}$	<6.0	$8.37^{+0.05}_{-0.04}$	$-0.43^{+0.09}_{-0.05}$	$0.17^{+0.18}_{-0.05}$	$-2.76^{+0.07}_{-0.05}$
SB 125	$7.06^{+0.08}_{-0.07}$	$-0.63^{+0.07}_{-0.08}$	$7.7^{+0.1}_{-0.1}$	<6.0	$8.38^{+0.03}_{-0.03}$	$-0.34^{+0.07}_{-0.07}$	$0.35^{+0.10}_{-0.14}$	$-2.58^{+0.06}_{-0.06}$
SB 126	$6.44^{+0.05}_{-0.11}$	$-1.42^{+0.05}_{-0.04}$	$7.9^{+0.0}_{-0.2}$	<6.0	$8.28^{+0.03}_{-0.03}$	$-0.52^{+0.03}_{-0.03}$	$0.15^{+0.06}_{-0.04}$	$-2.65^{+0.05}_{-0.06}$
SB 153	$4.58^{+0.04}_{-0.03}$	$-2.42^{+0.04}_{-0.03}$	$6.8^{+0.0}_{-0.0}$	<6.0	$8.36^{+0.03}_{-0.03}$	$-0.45^{+0.05}_{-0.04}$	$0.16^{+0.09}_{-0.04}$	$-2.51^{+0.07}_{-0.07}$
SB 179	$5.70^{+0.06}_{-0.06}$	$-1.62^{+0.05}_{-0.03}$	$7.3^{+0.1}_{-0.1}$	<6.0	$8.44^{+0.03}_{-0.03}$	$-0.35^{+0.04}_{-0.04}$	$0.17^{+0.08}_{-0.05}$	$-2.75^{+0.06}_{-0.06}$
SB 191	$4.51^{+0.05}_{-0.04}$	$-2.49^{+0.05}_{-0.04}$	$6.5^{+0.0}_{-0.0}$	<6.0	$8.51^{+0.02}_{-0.02}$	$-0.32^{+0.02}_{-0.02}$	$0.12^{+0.02}_{-0.01}$	$-2.09^{+0.05}_{-0.06}$
With UV stellar equivalent widths								
SB 9	$5.25^{+0.03}_{-0.03}$	$-1.87^{+0.02}_{-0.02}$	$7.1^{+0.0}_{-0.0}$	<6.0	$8.45^{+0.03}_{-0.03}$	$-0.22^{+0.02}_{-0.02}$	$0.46^{+0.03}_{-0.04}$	$-2.83^{+0.03}_{-0.03}$
SB 49	$6.17^{+0.04}_{-0.04}$	$-0.83^{+0.04}_{-0.04}$	$6.3^{+0.0}_{-0.0}$	$6.3^{+0.0}_{-0.0}$	$8.67^{+0.01}_{-0.00}$	$-0.00^{+0.00}_{-0.00}$	$0.49^{+0.01}_{-0.01}$	$-2.29^{+0.02}_{-0.01}$
SB 61	$7.33^{+0.08}_{-0.12}$	$0.08^{+0.04}_{-0.05}$	$7.3^{+0.1}_{-0.2}$	<6.0	$8.32^{+0.02}_{-0.04}$	$-0.46^{+0.02}_{-0.04}$	$0.22^{+0.03}_{-0.03}$	$-1.96^{+0.04}_{-0.04}$
SB 80	$6.09^{+0.04}_{-0.04}$	$-0.91^{+0.04}_{-0.04}$	$6.9^{+0.0}_{-0.0}$	<6.0	$8.33^{+0.03}_{-0.03}$	$-0.39^{+0.01}_{-0.01}$	$0.38^{+0.05}_{-0.06}$	$-2.59^{+0.04}_{-0.04}$
SB 119	$5.31^{+0.04}_{-0.04}$	$-1.69^{+0.04}_{-0.04}$	$6.8^{+0.0}_{-0.0}$	$6.1^{+0.0}_{-0.0}$	$8.37^{+0.04}_{-0.03}$	$-0.30^{+0.02}_{-0.01}$	$0.46^{+0.03}_{-0.06}$	$-2.56^{+0.04}_{-0.04}$
SB 125	$6.72^{+0.06}_{-0.04}$	$-0.33^{+0.03}_{-0.03}$	$7.0^{+0.1}_{-0.1}$	<6.0	$8.51^{+0.02}_{-0.01}$	$-0.13^{+0.02}_{-0.01}$	$0.49^{+0.01}_{-0.01}$	$-2.35^{+0.02}_{-0.02}$
SB 126	$6.17^{+0.04}_{-0.04}$	$-1.15^{+0.03}_{-0.03}$	$7.3^{+0.0}_{-0.0}$	<6.0	$8.48^{+0.01}_{-0.01}$	$-0.16^{+0.00}_{-0.00}$	$0.49^{+0.01}_{-0.01}$	$-2.40^{+0.02}_{-0.02}$
SB 153	$4.60^{+0.03}_{-0.02}$	$-2.40^{+0.03}_{-0.02}$	$6.7^{+0.0}_{-0.0}$	<6.0	$8.36^{+0.04}_{-0.04}$	$-0.37^{+0.01}_{-0.01}$	$0.35^{+0.07}_{-0.06}$	$-2.47^{+0.06}_{-0.06}$
SB 179	$5.41^{+0.03}_{-0.03}$	$-1.59^{+0.03}_{-0.03}$	$6.6^{+0.0}_{-0.0}$	$6.1^{+0.0}_{-0.0}$	$8.55^{+0.02}_{-0.01}$	$-0.09^{+0.01}_{-0.01}$	$0.49^{+0.01}_{-0.02}$	$-2.46^{+0.03}_{-0.03}$
SB 191	$4.54^{+0.03}_{-0.02}$	$-2.46^{+0.03}_{-0.02}$	$6.5^{+0.0}_{-0.0}$	<6.0	$8.55^{+0.00}_{-0.01}$	$-0.28^{+0.00}_{-0.00}$	$0.11^{+0.01}_{-0.01}$	$-2.01^{+0.03}_{-0.04}$

our objective is to determine whether the UV stellar wind lines can be reproduced, we also perform a second fit for each object including the equivalent widths of these features measured as for the data (as in Section 3.2, though obviously without the need for ISM or MW absorption line masking in the model spectra). For each object, we extract constraints on the underlying model parameters from the posterior (displayed in Table 6), as well as the flux distributions for the fitted lines to examine fit quality (Fig. 5). Finally, we also sample the posterior model distribution and extract predicted UV spectra at C IV $\lambda\lambda 1548, 1550$ and He II $\lambda 1640$. We fit a cubic spline to the UV spectrum with the stellar wind lines masked both for the data and the model spectra, and normalize the spectra by this continuum before comparison in Fig. 6.

Thus far and for the majority of this paper, we focus on fits to the optical nebular lines and the equivalent widths of the UV stellar wind lines. Comparison to the UV spectra only in a continuum-normalized space minimizes issues with reddening uncertainties and potential flux calibration systematics between the optical and UV. However, a mismatch in the stellar population probed by the SDSS fibre and slightly smaller COS aperture could still lead to systematic issues in reproducing these observables even in this simplified framework. In Appendix B, we analyse both the UV continuum and the He II stellar wind lines to argue that a large mismatch of this sort is highly unlikely. The stellar light captured by the *HST*/COS aperture appears consistent with having been produced by the same population of massive stars as is viewed in the SDSS fibre, indicating that aperture issues are unlikely to impact the results presented here.

While we focus on reproducing optical nebular line fluxes and continuum-normalized UV features to minimize systematics, we are also interested in whether the UV O III] and C III] lines can be

reproduced. The ratio of these lines also provides valuable constraints on the gas-phase C/O ratio. Thus, we perform an additional set of fits with the above-described optical line information along with the flux and equivalent width of both O III] $\lambda 1661, 1666$ and the total flux of the C III] $\lambda 1908$ doublet (with upper limits employed where lines are undetected or impacted by MW absorption), and allow the C/O abundance to vary over the full range of the models (0.1–1.4). We verify that the strength of the O III] and C III] lines are reasonably well-reproduced by this fit within the measurement uncertainties, confirming that the prominent C III] detections here can be reproduced by our population synthesis models and that the resulting C/O abundances are reasonable. We present these C/O constraints in Table 5. We also verify that the results of the experiment we present in the following section, comparing nebular line fits with or without including constraints on the UV stellar equivalent widths, produce qualitatively similar results for the stellar wind lines when the UV nebular lines are included. As already discussed above, for simplicity we focus on fits only to the optical nebular line and UV equivalent width fits for the remainder of this paper. However, we utilize the results of the joint optical and UV nebular line fits to first fix the C/O abundance assumed in the population synthesis fits presented and discussed here, adopting the median values presented in Table 5 for each.

The parameter estimates from the optical nebular line fits (Table 6) for the new systems in this paper are generally in reasonable agreement with the fits to the photometry and $W_0(\text{H}\beta)$ described in Section 3.1 (Table 3). The star formation rates are generally well-matched between the two methods within their uncertainty. The nebular line fits without UV information generally prefer a slightly earlier initiation of the current period of star formation than the photometry fits (median age of 20 Myr compared to 4 Myr), and the

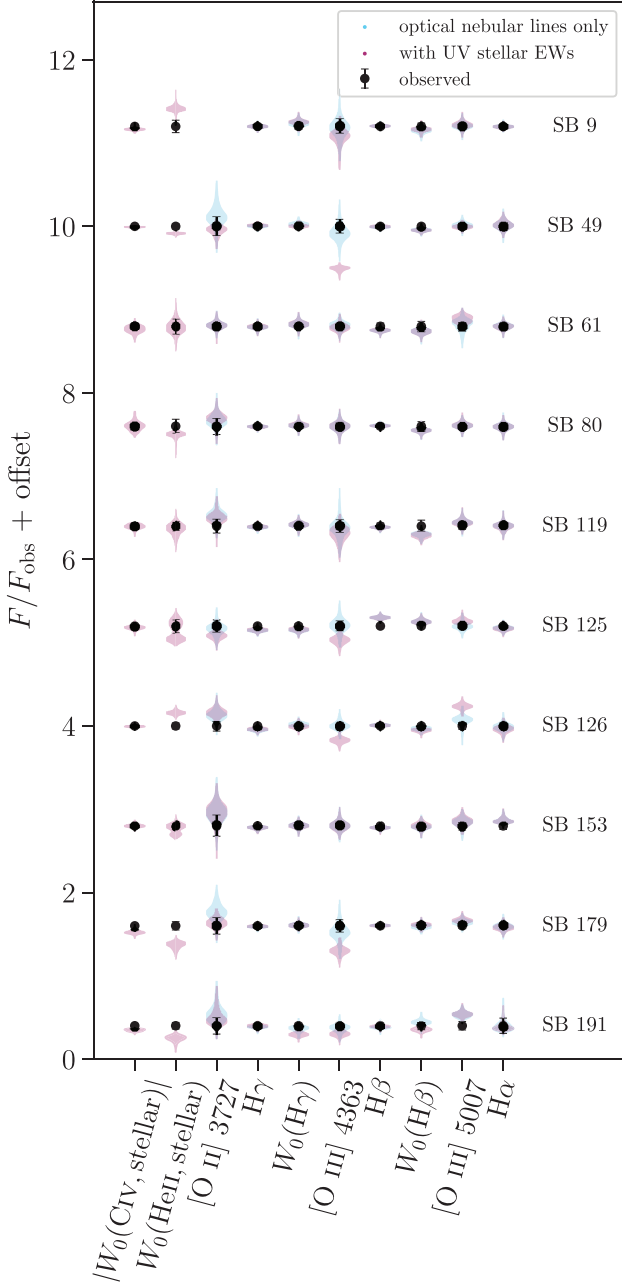


Figure 5. Results of our fits to nebular line measurements for each target system with BEAGLE. In each row, we plot the fitted measurement as a black point with error bar (often small enough to be hidden by the point); and behind it, we plot the fit model distribution for each normalized quantity drawn from the model posterior. We present results from two fits; one including only the optical nebular line information (presented in blue) and an otherwise identical run but incorporating the equivalent width of the UV stellar wind features presented on the right-hand side of this plot (red). Agreement is very good with fits to the nebular line information only, though some offsets with the data are apparent when jointly fitting these measurements with the stellar wind equivalent widths.

inferred stellar masses tend to be correspondingly higher by ~ 0.3 – 0.5 dex.⁷ The nebular-only fits uniformly prefer a truncation of the current period of star formation within the last Myr, and suggest gas-

⁷With the exception of SB 153, where the nebular line fit prefers a star formation rate 0.4 dex lower and a lower mass by 0.3 dex.

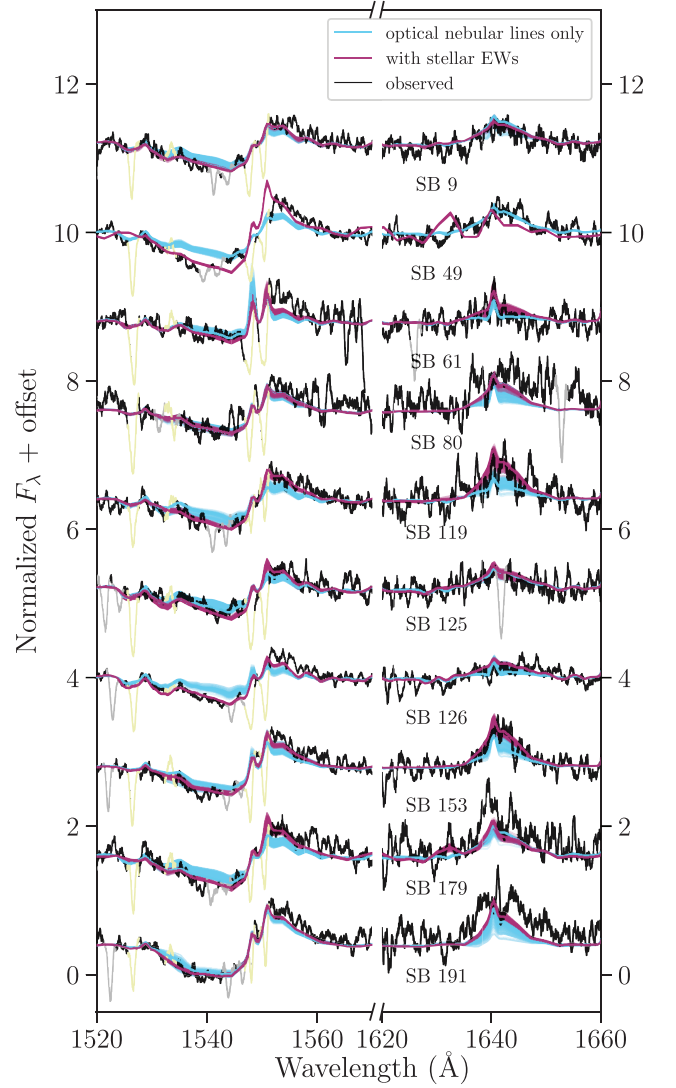


Figure 6. Comparison of our BEAGLE model results to the UV C IV and He II complexes, which in these galaxies are both dominated by broad massive stellar wind features. We plot both the data and models after first normalizing by the continuum as described in the text. The fits in blue incorporate no information about these wind features, whereas those in red include the equivalent width measured for both. Interstellar absorption lines in the data from the galaxy CGM and the MW are masked in yellow and white, respectively. Given only the nebular line constraints, the C&B models systematically underestimate the equivalent width of both stellar wind tracers. While incorporating these strengths into the fits generally improves the agreement, as discussed in the text this is in most cases requires metallicities and effective ages in significant tension with the optical data.

phase oxygen abundances $12 + \log \text{O}/\text{H} \simeq 8.3$ – 8.5 modestly larger than those inferred from the direct- T_e method (Table 4). With $\xi_d \simeq 0.1$ – 0.4 , this places the estimates of the stellar metallicity Z_* at 25–50 per cent solar. Adding the UV stellar equivalent widths changes this picture, as will be discussed in detail in the following section and Appendix C.

5.2 Results

In this subsection, we will analyse the results of the population synthesis fits described above and explore possible explanations for any discrepancies with the observations this uncovers. We will focus

first on the ability of the C&B models to reproduce the nebular lines and massive stellar wind signatures in 5.2.1. In 5.2.2 we will compare the metallicities derived using the C&B photoionization models and those computed with the direct T_e method (Section 3.1). Finally, we will discuss the peculiar profile of broad He II emission detected in two of the strongest emitters in our sample in 5.2.3.

5.2.1 Matching the strong UV stellar wind lines

In this section, we will summarize the joint analysis of the nebular emission lines and stellar wind signatures fit as-described in the previous section to determine whether the C&B models can self-consistently reproduce both. Matching both in local systems is a benchmark that stellar population synthesis prescriptions must meet if the constraints extracted from their application to nebular emission lines alone at high redshift are to be interpreted physically. While the C&B models we focus on do approximately reproduce the range of stellar wind strengths observed in our systems (Section 4), they do not include the impact of binary mass transfer or rotation. If these or other neglected physics play a significant role in shaping the evolution of the most massive stars, this may manifest in discrepancies with the spectroscopic properties of very young star-forming regions such as those presented in this paper. We defer a detailed case-by-case discussion of the objects to Appendix C, and summarize the results for the sample as a whole here.

The strong nebular emission lines powered by the massive stars in these systems provide leverage on both the gas-phase metallicity and ionization state as well on the stellar ionizing spectrum. In the first part of our experiment, we fit only the fluxes and equivalent widths of the optical nebular lines, and highlight the resulting model flux and equivalent width distributions relative to those observed in Fig. 5 (in blue). In every case, the models are able to match this gas emission well within the measured uncertainties. All of the fits prefer a continuous star formation history extending to the present (with the flexible star formation interval end uniformly constrained within $t_0 < 1$ Myr), but with very young maximum stellar ages ranging from 3–80 Myr (median 18 Myr). This is broadly congruent with the results of our broad-band SED fits including $H\beta$ equivalent widths (Table 3), though with a preference for somewhat older maximum ages and consequently slightly higher total stellar masses. The gas-phase metallicities inferred range from $12 + \log O/H = 8.25$ – 8.53 , systematically higher than those measured with the same lines using the direct- T_e method by ~ 0.2 dex on average (Table 4). We discuss this offset in more detail in Section 5.2.2, but conclude in summary that this is an expected consequence of comparing to estimates from photoionization models. The stellar metallicity Z is tied to this gas-phase metallicity in our model, with the caveat that increasing depletion on to dust grains with increasing dust-to-metal mass ratios (ξ_d) can allow for higher stellar Z at fixed gas-phase oxygen abundance (Section 5.1). But the models find good agreement with the nebular lines at ξ_d values uniformly below solar ($\xi_{d,\odot} = 0.36$; Gutkin et al. 2016), as expected for metal-poor star-forming galaxies; yielding stellar metallicities $Z/Z_\odot = 0.3$ – 0.5 . In all cases, the gas ionization parameter indicates highly ionized gas, with $\log U$ ranging from -3.0 up to -2.1 . In summary, the fits to the nebular lines alone paint a well-constrained picture of systems dominated by continuous star formation initiated recently with metallicities of 30–50 per cent solar.

As expected, our fits reveal that the properties of the stellar populations dominating these systems are specified fairly precisely by the strength and equivalent widths of the strong nebular emission

lines that their ionizing radiation powers, in the context of our stellar models. But the strong stellar wind lines accessed by our spectra provide an independent view of these massive stars against which these models can also be compared. In Fig. 6 we highlight the crucial C IV and He II complexes in the *HST*/COS data, and compare them to continuum-normalized model spectra sampled from the posterior of our fits to the nebular lines only (in blue). If the models have derived the correct mix of massive stars and treat their evolution and atmospheres correctly, we should expect to find good agreement with these wind lines.

Instead, the UV stellar wind lines reveal significant tension with the C&B model fits. In nearly every case, the strength of these features is underestimated by the models constrained by the nebular line properties; and there are no cases in which the model wind features are stronger than observed. The observations generally reveal more prominent broad He II emission as well as deeper absorption and stronger emission in the C IV P-Cygni profile than predicted. Both stellar wind diagnostics are systematically more prominent relative to the continuum than the C&B models predict given the nebular emission properties.

Before interpreting this result, we must determine whether there is any part of the model parameter space which would reconcile these differences with the observed stellar wind features. As described above, we repeat the same nebular line fits with the equivalent widths of the C IV absorption trough and He II emission line included alongside the optical line fluxes and equivalent widths. The resulting fits are displayed in red in Fig. 6, and as-expected, generally bring the UV stellar wind lines closer to agreement with the data. However, the only way that the models are able to approach the observed UV stellar wind equivalent widths is through changes in metallicity and age which bring the data into new tension with the optical data, as described in detail in Appendix C and summarized here.

First, in order to match the strong stellar wind profiles observed, the models in most cases attempt to move to higher stellar metallicities. Since the strength of line-driven stellar winds increases significantly with higher iron-peak element abundances, this allows the models to reach higher equivalent widths in both C IV and He II. This increase in the stellar iron abundance is accomplished for most systems by a combination of both an increase in the total metallicity Z (by up to 0.3 dex in the most extreme cases), as well as an increase in dust-to-metal mass ratio ξ_d (Table 6). As described above, since the gas-phase and stellar metallicity are tied, such an increase in depletion on to dust grains will lead to an increase in the stellar iron abundance if the gas-phase O/H is held constant. However, the preferred ξ_d is increased to above solar value in a total of seven of the ten systems, six of which reach the upper end of our prior distribution and the available models at $\xi_d = 0.5$. Such supersolar dust content is highly unlikely in these moderately low-metallicity star-forming regions. In addition, despite this shift in ξ_d , the increase in stellar Fe/H still induces a significant shift upwards in gas-phase O/H as well. In six of the ten cases, this shift upwards in gas-phase metallicity is large enough to bring the model confidence interval for the temperature-sensitive [O III] $\lambda 4363$ line flux out of agreement with the data due to enhanced cooling (Fig. 5). If we were to adopt a more flexible model with variable α/Fe which would effectively allow for decoupled stellar Fe/H and gas-phase O/H (e.g. Steidel et al. 2016), some of this tension could be resolved. However, achieving these results would require much higher iron content in the stars than suggested by the gas-phase O/H, implying a highly unlikely sub-solar α/Fe abundance (opposite that inferred for star-forming systems and theoretically expected from yields dominated by Type II supernovae).

In addition, the prominent UV stellar wind features force the models to different effective stellar age distributions. The C IV and He II profiles are dominated by initially very massive OIf and wind-stripped WR stars in the C&B models. As a result, the models for galaxies with initially underestimated UV wind features almost uniformly prefer a younger effective stellar ages, with the inferred start of star formation shifting down to a median of 7 Myr across all systems. In addition, the best fits for three systems with particularly discrepant UV wind features (SB 49, 119, 179) shifted from star formation proceeding continuously until $t_0 < 1$ Myr ago to preferring an early truncation 1–2 Myr ago; and in the case of SB 49, the preferred star formation history collapses to an essentially instantaneous burst occurring 2 Myr ago. However, these star formation history shifts result in very different populations of WR stars, and consequently different predictions for the optical WR wind features. For all three of these systems with the largest shift in star formation history, individual wind lines in the blue or red WR bumps are substantially overestimated relative to those observed (Fig. 7). This suggests that the improvements to the UV stellar wind lines for these systems are only accomplished in the models with populations of WR stars that are either present in the wrong numbers or at too high a metallicity, leading to significant disagreement in the optical.

Several other explanations also fall short of providing an adequate explanation for this difficulty in simultaneously matching the nebular lines and stellar wind features. It is unlikely these discrepancies are due to issues with our treatment of dust extinction or inter-aperture flux calibration, as we have focused only on the equivalent width of the stellar features in the UV (see also Appendix B). Our model included a flexible star-formation history allowed to vary from essentially instantaneous bursts to a continuous model; and adding an additional independent burst component allowed to vary in age on top of this model did not change the results qualitatively. While stochastic IMF sampling in very low-mass $\lesssim 10^5 M_\odot$ clusters could lead to peculiar populations of massive stars and spectroscopic properties disjoint from those of a continuous star-formation model (e.g. Paalvast & Brinchmann 2017; Vidal-García et al. 2017), the most discrepant objects are not biased towards particularly small numbers of massive stars. These systems extend to relatively large inferred masses $\sim 10^6 M_\odot$ (in particular, SB 49 and 179 at $10^{6.3}$ and $10^{5.7} M_\odot$, respectively) where stochastic sampling should be relatively unimportant. In particular, using a dust correction derived from the optical with an SMC extinction curve (Section 3.1) and assuming a single WNL star has a luminosity in He II $\lambda 1640$ of $1.2 \times 10^{37} \text{ erg s}^{-1}$ (Schaerer & Vacca 1998; Chandar et al. 2004), the integrated luminosity of He II in these systems suggests that the number of WR stars in the discrepant systems spans the entire range displayed by our sample; from as few as $N(\text{WNL}) \simeq 13$ in the lowest mass system SB 191 to as many as 300 and 1650 in SB 179 and 49, respectively. Thus, stochasticity is also unlikely to explain these offsets from the model predictions.

Rather than a product of analysis uncertainties, this systematic mismatch may be the signature of stellar physics that the C&B models do not presently include. The excess observed power in the UV stellar wind lines that the models cannot account for would be naturally produced by the massive, rejuvenated, and rapidly rotating products of binary mass transfer and mergers. Including such interactions especially at the high mass end would produce additional very massive stars with OIf or WN spectral signatures powered in their dense winds, as demonstrated by BPASS in the steady mass-transfer case (with a different set of underlying single-star evolutionary models with far weaker baseline stellar wind emission in He II).

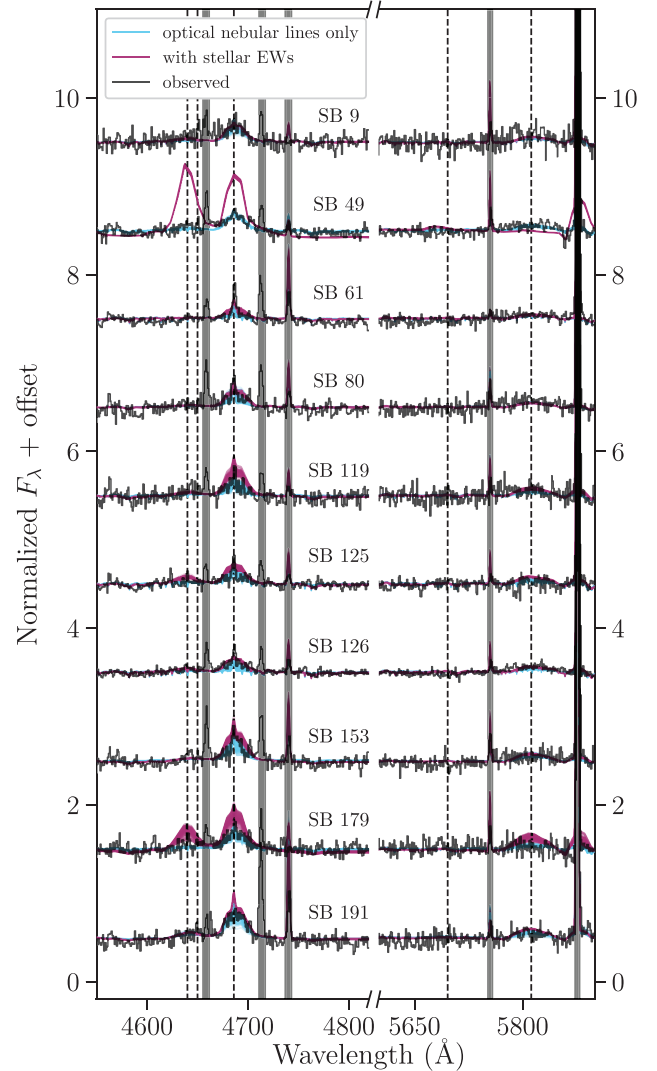


Figure 7. Comparison of our BEAGLE model results to the optical WR stellar wind features at 4600–4700 Å (the blue bump, left-hand side) and ~ 5650 –5800 Å (red bump, right-hand side). As in Fig. 6, we plot both the data and models respectively after first normalizing by the continuum. The fits in blue incorporate only nebular line information, whereas those in red also include the equivalent width of the C IV and He II stellar features measured in the UV. We highlight with dashed lines the approximate centres of the broad stellar wind features we aim to reproduce (N III 4640, C III/IV 4650, He II 4686 in the blue; and C IV 5808 in the red). The fits with the nebular line information-only provide a generally reasonable fit to the observed optical stellar wind features, in which He II is by far the most prominent; while the fits with the UV stellar wind strengths incorporated in some cases dramatically overpredict the strength of several of the WR stellar winds lines.

We discuss this possibility and its implications for galaxy modelling in more detail in Section 6.1.

5.2.2 Comparing gas-phase metallicity estimates

We note that fits to the optical nebular line information only consistently produce higher gas-phase metallicity estimates than found with the direct- T_e method above (Section 3.1). This offset ranges from 0.09–0.33 dex, with a median offset of 0.18 for our ten systems. This is a significant difference typically well in-excess of the joint uncertainties on both metallicity determinations. However,

comparable or larger offsets are common in the literature (see also Chevillard et al. 2018). The tendency of metallicity estimates based upon photoionization model grids to provide systematically larger gas-phase oxygen abundance estimates than the canonical two-zone direct- T_e method has been previously attributed to complications in interpreting [O III] $\lambda 4363$ with the oversimplified ionization structure assumed by the latter (e.g. Blanc et al. 2015; Vale Asari et al. 2016). Because the optical forbidden line ratios are sensitive to temperature and density fluctuations and to the velocity distribution of free electrons, the implicit assumption in the direct- T_e approach that [O III] $4363/5007$ provides an accurate bulk measure of the average T_e in a perfect two-zone H II region structure is likely not always valid (e.g. Tsamis et al. 2003; Nicholls, Dopita & Sutherland 2012). Indeed, this offset may be also related to the longstanding offset between metallicities determined with recombination lines and collisionally excited lines in nearby H II regions and planetary nebulae (the abundance discrepancy factor problem; e.g. Peimbert, Storey & Torres-Peimbert 1993; García-Rojas & Esteban 2007; Gómez-Llanos & Morisset 2020).

Regardless, since the lines including [O III] $\lambda 4363$ are well-fitted by BEAGLE when considered alone, we interpret this offset found when fitting the nebular lines only as the presently unmitigable systematic uncertainty in gas-phase oxygen abundance determinations using only collisionally excited nebular lines. In contrast, as discussed above, the far larger metallicity offsets found when the nebular lines and stellar wind features are fitted together are not well explained by these issues, and the inability to match the observed [O III] $\lambda 4363$ line fluxes is indicative of real tension with the data.

5.2.3 Peculiar He II profiles in SB 179 and 191

The objects which present the greatest disagreement with the stellar model predictions in the UV stellar wind lines are SB 179 and 191 (Section 5.2.1 and Appendix C). In particular, both power significantly more prominent He II $\lambda 1640$ emission than the models can reproduce alongside the optical nebular lines. Closer inspection of these very high-S/N stellar wind profiles in the UV could provide additional insight into the origin of the systematic disagreement of this line with the stellar models.

Intriguingly, both of these systems display a clear double-peak in the broad emission at He II $\lambda 1640$, with peak separation of order 1000 km s^{-1} (Fig. 8). No intervening MW absorption is expected at this wavelength for these systems (which reside at different redshifts, $z = 0.0047$ and 0.0028). If interpreted first as two kinematic emission components, this profile is inconsistent with a stellar origin. The two peaks are individually far more narrow than the other 1640 profiles, and 1000 km s^{-1} is well beyond the rotational velocity a massive star can maintain. Roughly symmetric double-peaked He II profiles are observed in some accretion discs surrounding neutron stars in low-mass X-ray binaries (e.g. Soria, Wu & Hunstead 2000; Val Baker, Norton & Quaintrell 2005) and in other nebular lines surrounding high-mass X-ray binaries (e.g. Filippenko et al. 1988). However, these systems typically display lower peak velocity separation ($\lesssim 500 \text{ km s}^{-1}$) than observed here. In addition, SB 191 has archival *Chandra* observations which constrain its X-ray emission to $L_{X, 0.5-8\text{keV}} < 1.3 \times 10^{38} \text{ erg s}^{-1}$, making the presence of a massive accretion disc in this system unlikely. Aspherical supernovae explosions are also observed to produce double-peaked emission line profiles with even larger velocity separations at late times in cases where the explosion is viewed close to down the jet axis (e.g.

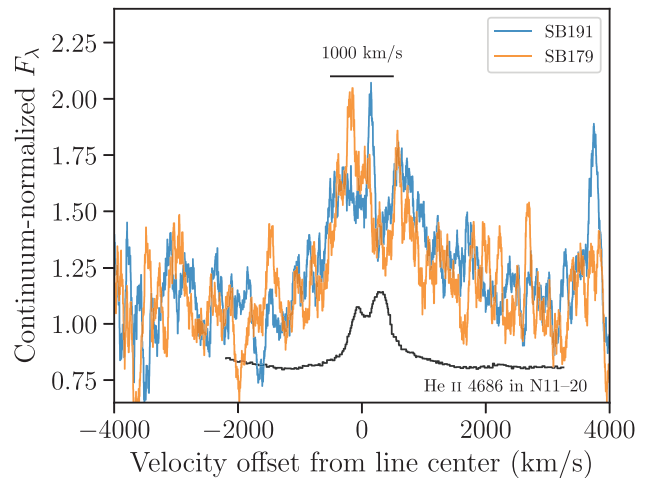


Figure 8. The peculiar, apparently double-peaked He II $\lambda 1640$ profiles for SB 179 and 191 (note SB 191 also shows a narrow nebular component in addition). For comparison, we plot the He II $\lambda 4686$ profile of the star N11–20, a rapidly rotating ($v \sin i = 260 \text{ km s}^{-1}$) Onfp giant in the LMC with the centrally absorbed profile characteristic of this spectral class (adapted from Walborn et al. 2010).

Maeda et al. 2008), but we see no other strong evidence for a recent supernovae in the nebular spectra.

Rather than a double-peak in emission, this profile may instead be characterized by central absorption. While this absorption appears too broad to be attributable to the underlying narrow photospheric profile normally observed in early-O dwarfs, such a centrally absorbed He II profile is actually observed in a rare subset of hot O star spectra, defining the spectral designation Onfp with He II $\lambda 4686$ in emission but with a peculiar central reversal (‘f’ and ‘p’; see Fig. 8; Walborn 1973; Conti & Leep 1974; Walborn et al. 2010). As indicated by the ‘n’, this profile is commonly found alongside significant broadening in the absorption and emission lines indicative of rapid rotation. Detailed atmosphere modelling has demonstrated that such a profile is naturally produced for rotating stars driving an isotropic wind (Hillier et al. 2012), supporting the notion that such stars are rapid rotators driving dense stellar winds. Such strong winds should efficiently brake rotation, but many of the stars of this class in both the Milky Way and Magellanic Clouds show evidence of binarity or runaway velocities (Walborn et al. 2010). They may then be the product of binary mass transfer or stellar mergers, rejuvenated in luminosity and spun-up to very high velocity by this interaction.

While a promising potential clue as to the nature of the dominant massive stars in these systems, several aspects of the observed integrated He II profiles remain puzzling in this context. The He II profiles for individual Onfp stars typically show slightly blueshifted absorption, and thus an asymmetric emission profile with the red component brighter. However, this effect may be washed-out when integrated over a population of stars at a spread in velocities powering various He II profiles. In addition, the optical He II 4686 profiles for SB 179 and 191 do not show the same morphology, even at Keck/ESI resolution (S17). Individual Onfp stars show significant time variability in their wind profiles on time-scales of days however, sometimes transitioning out of this peculiar profile entirely (Walborn et al. 2010). We demonstrate in Appendix B that the He II 1640/4686 ratio for our targets is consistent with there being no severe mismatch in the stellar population probed by the SDSS and *HST*/COS apertures. However, different effective dust attenuations for nebular gas and various populations of stars may still impact the integrated profiles

of these lines. This may be the part of the cause of the puzzling lack of a nebular He II λ 1640 detection in the *HST*/COS spectrum of SB 125 and 126 when narrow emission is present at He II λ 4686 in the SDSS spectrum. And if Onfp stars are indeed likely runaways expelled from their natal clusters by binary interaction, they may see a substantially smaller column of dust than other massive stars and thus are more likely to dominate the He II profile in the ultraviolet. High spatial and spectral resolution optical IFU observations probing the He II λ 4686 line across these systems are necessary to confidently confirm and narrow the origins of this peculiar emission profile. However, this detection may lend unexpected support to the notion that the excess emission in He II is indeed due at least in part to rapidly rotating products of mass transfer, as discussed further in Section 6.1.

6 DISCUSSION

We have presented new ultraviolet spectra for a sample of nearby star-forming regions hosting very young stellar populations at moderately subsolar (~ 20 – 50 per cent Z_{\odot}) metallicity. Such integrated spectra provide a key opportunity to study the properties of the massive star populations that dominate these clusters and similarly high specific star formation rate systems at cosmological redshifts. Joint fits to both the strong stellar wind emission signatures and the nebular emission the massive stars power reveal significant tension with the data, suggestive of potentially missing physics in the stellar population synthesis models at these very young ages. We also uniformly detect nebular C III] emission in this sample, enabling an investigation of the strength of this doublet for galaxies dominated by very recent star formation. We first discuss the implications for models of very young stellar populations and massive stars in Section 6.1, before turning our attention to interpreting strong C III] emission at high-redshift in Section 6.2.

6.1 Evidence for an overabundance of very massive stars

Nebular spectroscopy at cosmological distances, soon to be extended to large samples at $z > 6$ with *JWST*, represents an opportunity to study early star-forming systems in much greater physical detail than photometry alone can afford. However, the accuracy of physical inferences drawn from fitting this emission will be fundamentally limited by the uncertainties in stellar population synthesis models, which still vary substantially with different prescriptions for stellar evolution especially at ionizing energies (e.g. Stanway & Eldridge 2019). Evidence from both photometric bands contaminated by rest-optical line emission and the first rest-UV spectra in the reionization era suggests that a significant fraction of $z > 6$ galaxies are dominated by populations of massive stars formed in bursts initiated within < 20 Myr of observation (e.g. Stark et al. 2015b; Endsley et al. 2021). Thus, our understanding of the most distant galaxies *JWST* will observe will depend sensitively on models for the evolution of populations of massive stars on much shorter time-scales than typical star-forming galaxies even at $z \sim 2$ (e.g. Steidel et al. 2016).

Study of stars and clusters in the Local Group has provided important clues as to the physics shaping the evolution of very young stellar populations. In particular, the 30 Doradus star-forming region in the LMC and its central young cluster R136 (~ 1 – 3 Myr, $Z/Z_{\odot} \sim 0.5$) has long represented a benchmark analogue for intense starbursts observed at high redshift. Intriguingly, the resolved stellar contents of this regions reveal evidence for a significant overabundance of massive stars $> 30 M_{\odot}$ relative to a Salpeter IMF (Schneider et al. 2018), including a sizable population of very massive stars with initial masses 100 – $320 M_{\odot}$ in R136 (Crowther et al. 2016). Simulations

suggest that such an overpopulation can be readily produced by mass transfer and mergers in binary systems, yielding numerous very massive stars even exceeding the upper limit of the initial mass function on time-scales of Myr (e.g. de Mink et al. 2014; Schneider et al. 2014). These very massive OIf/WN stars drive dense winds yielding prominent He II and C IV (Crowther et al. 2016) and, if they acquire sufficient angular momentum during mass transfer, may undergo quasi-chemically homogeneous evolution and effectively become an extremely massive Wolf–Rayet star. The stars produced by this mechanism are very good candidates for long gamma-ray burst progenitors (e.g. Cantiello et al. 2007), and may have an outsized impact on the properties of recently formed star clusters. However, without additional resolved star-forming region approaching the age and luminosity of 30 Doradus (e.g. Kennicutt 1984), the prevalence of these extreme mass transfer products and their impact on the integrated spectra of young star-forming regions as a function of metallicity remains unclear.

While unresolved, young star-forming regions such as those presented in this work provide the next best laboratories in which the evolution of populations of massive stars can be constrained. The systems presented in this paper are by-selection dominated by very young stellar populations, with [O III]+H β equivalent widths and photometric sSFRs at or exceeding the highest values inferred at $z > 6$ and suggesting these systems provide a clear window on to very recently formed $\lesssim 10$ Myr stellar populations (Tables 3 and 4). And also by selection, all reside at sufficiently high metallicities ($Z/Z_{\odot} \sim 20$ – 50 per cent) for radiatively driven stellar wind signatures to be detectable in the integrated spectra alongside strong ionized gas emission. The He II λ 1640 stellar wind line is a particularly useful probe of both initially very massive and luminous massive stars rejuvenated by mass transfer (see also Stanway et al. 2020). Reproducing this stellar wind line along with the other available constraints is a crucial litmus test for stellar population synthesis, but matching it at high equivalent width in both local and $z \sim 2$ galaxies has long stymied stellar models (Section 4).

In Section 5, we described an experiment designed to leverage both the nebular and stellar wind lines to search for evidence of stars missing from current stellar population synthesis models. We focus on the C&B stellar population synthesis models, which adopt the latest prescriptions for single-star evolution but do not include binary mass transfer or rotation. While BPASS incorporates a model for the impact of mass transfer on very massive stars, the predicted impact on He II on < 10 Myr time-scales in this framework is minimal, and insufficient to reach the equivalent widths observed in our local young star-forming regions (Section 4 and Eldridge & Stanway 2012). Significant recent improvements in wind-driven mass-loss predictions have allowed the C&B models to resolve much of the tension with the strength of He II λ 1640 relative to older stellar population models (Section 4), so residual issues in matching this line may be a relatively clean probe of additional physics these models are missing at very young ages.

Our comparisons with the C&B models revealed significant tension with the observed star-forming regions. When the nebular lines alone were fit, we found that the models systematically underestimated the strength of both the UV stellar C IV P-Cygni and He II emission lines (Section 5.2.1). Explicitly including the UV stellar equivalent widths in the fits failed to identify a consistent model able to match both the nebular line measurements and the UV–optical stellar wind strengths, suggesting that the models would require unreasonably high metallicity stars to reach the observed wind strengths. We discussed other possible explanations for this offset, but find no satisfactory alternative explanation, suggesting

that even with their improved treatment of massive stellar winds, the C&B models simply cannot self-consistently reproduce the nebular emission and observed wind line equivalent widths.

The mismatch between the C&B predictions and our data in the UV stellar wind lines is consistent with a significant overabundance of very massive stars relative to the models. In R136, stars in excess of $100 M_{\odot}$ completely dominate He II and contribute strongly to C IV due to the particularly strong winds these luminous stars drive (Crowther et al. 2016). A top-heavy initial mass function could explain this, though the existing evidence for significant IMF variations at subsolar metallicities is not clear-cut (e.g. Gunawardhana et al. 2011; Marks et al. 2012; Hopkins 2018); but as outlined above, a high incidence of mass transfer and mergers among initially very massive stars would also naturally yield such a distribution. If binary evolution has substantially affected the massive stars present, we would expect many to be rapidly rotating due to angular momentum transfer. The detection of peculiar double-peaked profiles which may be produced by rapidly rotating stars in two of the most discrepant systems presented lends additional support to the notion that mass transfer occurring on very short time-scales may be at work (Section 5.2.3).

Our results suggest that the stellar populations in very young systems with sSFR comparable to bursts observed at $z \sim 6$ may host populations of very massive stars significantly enhanced by mass transfer and mergers. Much of the work on the observational consequences of binary evolution for the SEDs of star-forming galaxies has focused on the impact of these processes on ~ 10 – 100 Myr time-scales when numerous initially lower mass stars ($\lesssim 20 M_{\odot}$) become significantly hotter than canonically expected (e.g. Eldridge & Stanway 2012; Ma et al. 2016; Xiao et al. 2018). However, our difficulty in matching the ultraviolet spectra of these young star-forming regions suggests that binary evolution may profoundly alter the properties and relative number of very massive stars on significantly shorter time-scales as well. The fact that this results in detectable differences in the integrated spectra of young star-forming regions has dramatic consequences for attempts to model systems dominated by similar populations at high-redshift.

These data indicate that neither C&B nor BPASS can fully describe the properties of young stellar populations dominated by very massive stars. This suggests that the mapping from metallicity and age to stellar population composition and the emergent ionizing spectrum may be significantly in-error at ages < 10 Myr. At yet lower metallicities (< 20 per cent Z_{\odot}) where stellar winds weaken and become increasingly transparent to hard ionizing radiation, the same stellar products of binary evolution will be far more difficult to observe in stellar wind emission but (for exactly the same reason) will likely have an even more dramatic impact on the integrated ionizing spectrum and thus observed nebular emission. It is crucial that these models be carefully tested at these young ages to mitigate the impact of these systematic uncertainties on our ability to understand the spectra of high-sSFR systems in the early Universe.

The nebular emission and rich continuum signatures accessible simultaneously in extreme local objects such as these likely represent one of our best opportunities to directly constrain models for the evolution of populations of very massive stars. Combined with observations of additional UV wind lines which can provide more detailed leverage on the detailed makeup and age of the massive star population (e.g. Smith et al. 2016; Chisholm et al. 2019), self-consistently reproducing UV–optical spectra for young star-forming regions such as those presented here will be an important benchmark for next-generation stellar population synthesis models incorporating the substantial improvements made both to prescriptions for single and binary star evolution over the past decade. Successfully matching

the stellar and nebular features in these systems would be strong evidence that the mapping between stellar metallicity, age, and ionizing spectrum predicted by the models is accurate at very young ages, directly addressing one of the chief sources of systematic uncertainty in the modelling of high-redshift galaxies.

6.2 Nebular C III] emission powered by moderately metal-poor massive stars

The growing body of rest-frame UV line detections at $z \gtrsim 6$ has ignited substantial interest in locating and understanding nearby galaxies that power similar emission. Particularly puzzling among the four C III] detections at these redshifts are those in EGS-zs8-1 and z7_GND_42912, two very bright $L_{UV}^* \simeq 3$, 2 reionization-era galaxies with C III] measured at 22, 16 Å (respectively; Stark et al. 2015a; Hutchison et al. 2019). This is an order of magnitude higher than typical galaxies at similar luminosities at lower redshift (e.g. Shapley et al. 2003; Du et al. 2017). The emerging picture from extensive work at both $z \sim 2$ – 4 (e.g. Nakajima et al. 2018; Mainali et al. 2020; Du et al. 2020) and in the local $z \sim 0$ Universe (e.g. Rigby et al. 2015; Berg et al. 2016; Peña-Guerrero et al. 2017; Senchyna et al. 2017, 2019) suggests that such prominent C III] emission requires a combination of both very high specific star formation rates $\gtrsim 10$ – 100 Gyr^{-1} and very metal-poor ($Z/Z_{\odot} \lesssim 0.2$, $12 + \log O/H \lesssim 8.0$) stars and gas. If such low metallicities are indeed necessary, these prominent C III] detections at $z > 6$ imply that these $\gtrsim L^*$ systems reside at sub-SMC metallicities, implying a large shift in the luminosity–metallicity relationship from $z \sim 2$ – 6 .

However, the most extreme of the local galaxies with UV spectroscopy hint at a possible alternate solution. The detection of very strong ~ 9 – 11 Å C III] emission alongside prominent stellar He II in SB 179 and 191, very young local star-forming regions at $12 + \log O/H = 8.3$ ($Z/Z_{\odot} \simeq 0.4$), suggested that very massive stars could potentially power stronger emission than expected at well above SMC metallicities (S17). The spectra presented in this work for seven additional systems selected in the same manner as SB 179 and 191 allow us to test this possibility. In addition to their WR wind signatures in the optical, our targets extend to very high sSFRs, and reach equivalent widths in [O III] $\lambda 5007$ up to 1719 ± 133 Å (in SB 153) higher than in any of the targets studied in S17. According to the trend towards more prominent C III] emission at high sSFR or $W_0([O III])$ displayed by that sample and observed by others (e.g. Mainali et al. 2020), such young effective stellar populations should be capable of powering equivalent widths in C III] approaching that observed in the reionization era.

In Fig. 9 we plot the observed equivalent width of C III] emission in these systems and other UV observations of $z \lesssim 0.3$ star-forming galaxies. Despite their extraordinarily high equivalent width optical emission lines and correspondingly young effective stellar ages, none of these objects power C III] emission above 12 Å. Even SB 153, with extremely prominent wind emission at 3.9 ± 0.1 Å in He II $\lambda 1640$ and an effective young stellar population age of $\lesssim 5$ Myr (assuming a recent CSFH), only reaches 8.7 ± 0.4 Å in C III]. This suggests that we are observing the upper envelope of the expected strength of C III] at metallicities $12 + \log O/H > 8.0$.

Besides metallicity and sSFR, several other factors play a role in regulating C III] emission. The ISM abundance of carbon has a direct effect on the strength of C III], and can suppress this emission significantly at very low C/O. However, our galaxies reside at moderate C/O abundances ($-0.7 < \log_{10}(C/O) < -0.2$, with median -0.31 ; Table 5), consistent with other local systems at comparable oxygen abundances but significantly higher than the bulk of more

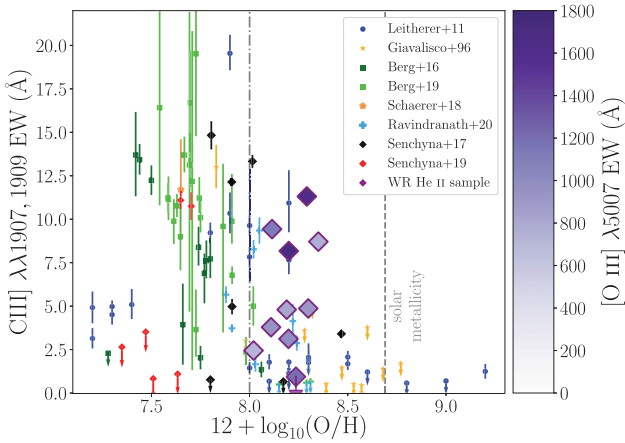


Figure 9. The equivalent width of nebular C III] emission as a function of metallicity for the ten extreme WR galaxies presented in this work. We also show measurements from UV spectra of $z \lesssim 0.3$ galaxies assembled by Leitherer et al. (2011), Giavalisco, Koratkar & Calzetti (1996), Berg et al. (2016), Senchyna et al. (2017, 2019), Schaeerer et al. (2018), Berg et al. (2019a), and Ravindranath et al. (2020). Oxygen abundances corresponding to solar and 20 per cent solar (SMC metallicity) are indicated by vertical dashed lines. While the galaxies presented here extend to extremely high sSFR (indicated by their [O III] $\lambda 5007$ equivalent widths) and all but one are detected in C III], they all remain below 12 \AA in equivalent width in this doublet. This suggests that even galaxies dominated by very massive stars cannot power emission in C III] at the level observed at $z > 6$ at the moderately low metallicities of our sample ($Z/Z_{\odot} \sim 0.2\text{--}0.5$).

metal-poor galaxies at $12 + \log O/H < 8$ which extend down to $\log_{10} C/O = -1.0$ (e.g. Esteban et al. 2014; Berg et al. 2019a). Finding star-forming galaxies at high redshift with substantially higher C/O than these systems is unlikely given this observed pseudo-secondary behaviour and low-metallicity plateau in C/O in local systems. We do expect to encounter α/Fe enhancement in high-redshift galaxies dominated by a rising star formation history, which can lead to a lower metallicity stellar population and thus harder ionizing radiation field than expected from the gas-phase oxygen abundance (e.g. Steidel et al. 2016; Strom et al. 2018; Sanders et al. 2020). The presence of lower metallicity stars may act to further enhance C III] even at moderate gas-phase oxygen abundance, though the magnitude of the effect will depend on the interplay between both gas ionization shaped by these stars and cooling efficiency determined largely by O/H.

Even for systems dominated by very young stellar populations with prominent optical [O III] emission, it appears unlikely that C III] significantly in-excess of 10 \AA can be powered for stars and gas at metallicities above $12 + \log O/H > 8.0$ ($Z/Z_{\odot} \gtrsim 0.2$). The number of systems at very high sSFR studied at metallicities $12 + \log O/H > 8.4$ remains relatively small (Fig. 9), but all are consistent with even weaker C III] $\lesssim 5 \text{ \AA}$. This is broadly consistent with the picture painted by photoionization models (e.g. Jaskot & Ravindranath 2016; Nakajima et al. 2018; Plat et al. 2019), which indicate that both softer ionization radiation fields and lower electron temperatures encountered at higher metallicities conspire to suppress collisionally excited emission from the C III] doublet. While C III] and O III] have similar ionization potentials and our systems power extremely prominent [O III] emission, the C III] doublet is far more sensitive to electron temperature than the optical [O III] lines (e.g. Jaskot & Ravindranath 2016; Tang et al. 2021). As a result, C III] is suppressed more strongly by the drop in temperatures associated with the increased efficiency of metal-line cooling at these high gas-

phase metal abundances, producing the relatively sharp evolution with metallicity we observe. The possibility remains that some of the strongest emission $\gtrsim 20 \text{ \AA}$ encountered in this line at high-redshift may require additional photoionization and heating from AGN or fast radiative shocks (Nakajima et al. 2018; Le Fèvre et al. 2019; Plat et al. 2019). However, in galaxies with nebular line ratios consistent with star formation, detection of C III] emission at equivalent widths in-excess of 15 \AA can be considered strong positive evidence for metallicities below that of the SMC ($12 + \log O/H \lesssim 8.0$). Along with C IV and He II emission which become prominent at yet lower metallicities ($12 + \log O/H < 7.7$; e.g. Berg et al. 2019b; Senchyna et al. 2019), these strong UV nebular lines may be a powerful probe of early chemical evolution in the *JWST* era when detection of [O III] $\lambda 4363$ is intractable due to faintness or when the optical lines are entirely inaccessible at the highest redshifts.

7 SUMMARY

Nearby star-forming regions at very young ages represent a crucial testbed for models of massive stellar populations and a benchmark for understanding high-redshift galaxies. In this paper, we presented new ultraviolet spectra probing the highly ionized gas and stellar winds in seven systems (combined with an additional three systems originally presented in Senchyna et al. 2017) which are all dominated by very young stellar populations identified via high equivalent width nebular emission alongside WR wind signatures in the optical (Section 2). The optical and ultraviolet spectra both reveal clear signatures of massive stars at gas-phase metallicities corresponding to 25–40 per cent solar (Section 3), and effective stellar population ages $\lesssim 10$ Myr comparable to that inferred for systems undergoing intense bursts of star formation at $z > 6$. Our spectra access both C III] and O III] nebular emission as well as stellar He II $\lambda 1640$ emission and C IV P-Cygni features produced in dense stellar winds (Section 4). These high-quality spectra enable a unique test of stellar population synthesis models, leveraging constraints on both the ionizing spectrum and metallicity through nebular gas emission and on the strong winds driven by massive stars through the wind diagnostic lines. Our main conclusions are summarized as follows:

(i) We detect nebular C III] in all of the targeted systems, confirming its ubiquity at high sSFR in galaxies below 50 per cent solar metallicity. However, the equivalent widths in this doublet only reach a maximum of 12 \AA for the entire sample of local galaxies in this metallicity range ($8.1 < 12 + \log O/H < 8.4$), likely due to efficient gas cooling at these metallicities. This suggests that the most prominent emission powered by stars in this line ($> 15 \text{ \AA}$) requires gas and/or stars at or below SMC metallicity ($Z/Z_{\odot} \lesssim 20$ per cent), suggesting that the massive reionization-era systems detected at these equivalent widths may harbour significantly subsolar gas and stars.

(ii) We detect He II $\lambda 1640$ stellar wind emission at equivalent widths ranging up to 4.7 \AA , among the most prominent detections of this line known and well in-excess of previous generations of stellar population synthesis predictions. We find that the latest Charlot & Bruzual (C&B, in preparation) models are able to largely reproduce this observed range, while the current predictions from BPASS are not. This suggests that improvements in the treatment of massive star evolution and spectra in the C&B models may play an important role in reproducing the spectra of the most extreme star-forming regions both locally and at $z > 6$.

(iii) However, we are unable to simultaneously match both the strong UV stellar wind signatures and the nebular line measurements with the C&B stellar population models. These models, which

neglect binary mass transfer, systematically underpredict the strength of both C IV and He II in the UV. Any C&B solution able to match these wind strengths resides at too large a metallicity, bringing the optical nebular lines and WR features out of alignment with the data. We conclude that the most plausible explanation for this discrepancy is an overabundance of the very massive stars that dominate these wind lines at young ages, which would be naturally expected from a high incidence of binary mass transfer and mergers. This suggests that binary evolution may substantially alter the properties and appearance of populations of massive stars on very short time-scales <10 Myr to which observations of high sSFR galaxies are sensitive.

As stellar population synthesis models become increasingly sophisticated and their applications to observables at high-redshift become increasingly varied, it is ever more important to anchor their predictions empirically where possible. Our results suggest that current prescriptions for massive stellar populations fall short of reproducing the properties of the highest sSFR systems nearby at ~ 30 per cent solar metallicity. Resolving this discrepancy will likely involve synthesizing the progress made by different modelling groups in treating the evolution and properties of binary mass transfer products and very massive stars. The high-quality spectra attainable for nearby star-forming systems such as these represent one of our only opportunities to test this next generation of stellar population models as a function of age and metallicity, with important consequences both for our understanding of massive star evolution and for correctly modelling the very young galaxies in the early Universe that *JWST* is poised to uncover.

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ern Reserve University, University of Chicago, Drexel University, Fermilab, the Institute for Advanced Study, the Japan Participation Group, Johns Hopkins University, the Joint Institute for Nuclear Astrophysics, the Kavli Institute for Particle Astrophysics and Cosmology, the Korean Scientist Group, the Chinese Academy of Sciences (LAMOST), Los Alamos National Laboratory, the Max-Planck-Institute for Astronomy (MPIA), the Max-Planck-Institute for Astrophysics (MPA), New Mexico State University, Ohio State University, University of Pittsburgh, University of Portsmouth, Princeton University, the United States Naval Observatory, and the University of Washington.

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DATA AVAILABILITY

The *HST* and SDSS data underlying this article are available through their respective data repositories (see <https://archive.stsci.edu/> and <https://www.sdss.org/dr16/>). Data will also be shared upon reasonable request to the corresponding author.

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APPENDIX A: TARGET STAR-FORMING REGIONS IN CONTEXT

Here we summarize the target objects in the context of their host galaxies and relevant previous studies in which they have been identified.

SB 9 (SDSS J130432.27-033322.0, Plate-MJD-Fibre 339-51692-83) is a star-forming region embedded in the northern outskirts of the barred-spiral galaxy UGCA 322. This system was also identified as WR 33 in the catalogue of SDSS spectra with prominent optical Wolf-Rayet features presented by Brinchmann et al. (2008b).

SB 49 (SDSS J011533.82-005131.1, Plate-MJD-Fibre 695-52202-137, WR 354) is located in an eastern arm of the spiral galaxy NGC 450.

SB 61 (SDSS J144805.38-011057.6, Plate-MJD-Fibre 308-51662-81,⁸ WR 25/182) is an isolated compact galaxy at $z = 0.0274$, placing it at the greatest estimated distance for our sample (117 Mpc).

SB 119 (SDSS J143248.36+095257.1, Plate-MJD-Fibre 1709-53533-215, WR 380) is a star-forming region in the southeastern outskirts of the barred spiral NGC 5669.

SB 125 (SDSS J113235.34+141129.8, Plate-MJD-Fibre 1754-53385-151, WR 384) is a lone compact star-forming region at the northeastern end of the tadpole-like spiral galaxy IC 2919.

SB 126 (SDSS J115002.73+150123.4, Plate-MJD-Fibre 1761-53376-636, WR 385) is an isolated compact star-forming galaxy also identified as Mrk 750.

SB 153 (SDSS J131447.36+345259.7, Plate-MJD-Fibre 2023-53851-263, WR 473) is the brightest star-forming region in the irregular galaxy Mrk 450.

APPENDIX B: ON COMPARING SDSS AND HST/COS SPECTRA

As described in Section 5.1, we have intentionally chosen to focus primarily on the equivalent width rather than the absolute flux of stellar features in the UV spectra to minimize the impact of uncertainties in reddening and flux calibration from the optical to UV. However, a systematic mismatch in the population of stars probed by the SDSS and COS apertures would still confound such an analysis. Given the typical pointing uncertainty of *HST* and the successful acquisition of a bright compact central source in all visits

(Fig. 2), it is very likely that the *HST*/COS aperture is centred on the same star-forming complex as SDSS in all cases. However, it is plausible that the smaller 2.5 arcsec diameter *HST*/COS PSA might miss some stellar flux captured by the 3 arcsec SDSS fibre. If this was the case, we would find systematically lower fluxes both in the UV continuum and stellar wind lines relative to that expected from the optical spectra. In this subsection, we describe the results of two independent experiments which together provide strong evidence that such a mismatch has not taken place.

First, we consider the optical SED fits described in Section 3.1. These BEAGLE fits were performed with only the SDSS fibre-matched *ugriz* magnitudes and $H\beta$ equivalent widths, with no information about the UV. If the *HST*/COS PSA was missing flux or offset from the SDSS aperture, we would expect these optical fits to systematically overpredict the flux measured at FUV wavelengths. To test this, we examine the full SEDs predicted by BEAGLE across the posterior distribution of these fits. For each galaxy, we generate a model-predicted median FUV flux weighted by the model likelihood in the relatively clean interval of 1560–1630 Å between the C IV and He II UV wind lines, which naturally includes both the model stellar and nebular continuum contributions. Comparing this to the observed median *HST*/COS G160M flux in the same interval reveals prediction-normalized offsets with median and standard deviation of (observed – model) = -12 ± 23 per cent. Thus, we find no statistically significant systematic offset between the *HST*/COS data and UV model expectations derived from the optical, even with no constraints provided on the UV–optical slope or dust attenuation, making it unlikely that the optical fibre has captured significantly more massive stars than the UV aperture.

And secondly, we directly leverage the strong stellar He II wind signatures present in both the optical and UV spectra. Because both He II $\lambda 1640$ and $\lambda 4686$ are recombination line complexes, their intrinsic ratio at fixed gas temperature and density under Case B conditions can be well-estimated theoretically (Hummer & Storey 1987). Indeed, this appears to hold even under the extreme conditions of Wolf-Rayet stellar atmospheres: Leitherer, Lee & Faisst (2019) demonstrate observationally that this line ratio is tightly constrained to a common value for a wide range of WN stars in the Milky Way and Magellanic Clouds. Thus, analogous to the Balmer decrement, the stellar He II 1640/4686 ratio can be used to estimate the stellar reddening. If the *HST*/COS aperture enclosed only a subset of the WR stars captured by SDSS, we would expect to find a systematically depressed He II 1640/4686 ratio, mimicking the effect of enhanced reddening. While the He II $\lambda 4686$ line is detected at relatively low signal-to-noise in SDSS relative to the $\lambda 1640$ profiles observed with *HST*/COS, we find that a reasonable constraint is obtained by simply integrating over a broad ± 20 Å window with a local continuum determined after masking the nearby nebular lines [Fe III] $\lambda 4658$, [Ar IV]+He I $\lambda 4712$, and [Ar IV] $\lambda 4740$. Directly comparing this line flux with that measured similarly for He II $\lambda 1640$ (Table 5) reveals raw ratios of 1640/4686 ranging from 3.2–10.6 (see Table B1). This maximum is larger than the median intrinsic ratio estimated by Leitherer et al. (2019) but well within the scatter they observe for individual WN stars. When translated into estimates of $E(B-V)$ via the relations from their table 5 after correction for Galactic extinction (as in Section 3.1), these ratios correspond to a median (maximum) in $E(B-V)$ of 0.02 (0.11) assuming an SMC extinction curve (Gordon et al. 2003, as in Section 3.1) and 0.03 (0.21) assuming a shallower Milky Way relation (Mathis 1990). That is, the inferred reddening towards the massive stars in the target systems as a whole is generally similar to or smaller than that inferred from the nebular Balmer decrement (Table 4). This is broadly consistent with the

⁸We use this SDSS spectrum over the other flux-matched spectrum 920-52411-575 as it is identified as the sciencePrimary deliverable by SDSS.

Table B1. Integrated flux ratio of the He II $\lambda 1640$ Å and 4686 Å recombination lines measured respectively in the *HST/COS* and SDSS spectra for our targets. We include estimates of the reddening inferred from this ratio after correction for Galactic extinction using the relations published by Leitherer et al. (2019) adopting both an SMC and MW extinction curve. The resulting range of reddening estimates is comparable to that inferred from the Balmer decrement (Table 4), suggesting that the *HST/COS* aperture probes the same population of massive stars captured by SDSS.

Name	$f(\lambda 1640)/f(\lambda 4686)$	$E(B-V)$	
		SMC	MW
SB 9	5.97 ± 1.22	0.02	0.04
SB 49	6.31 ± 0.68	0.01	0.02
SB 61	4.41 ± 0.64	0.05	0.10
SB 80	3.20 ± 0.57	0.11	0.21
SB 119	10.58 ± 2.25	0.00	0.00
SB 125	10.55 ± 1.96	0.00	0.00
SB 126	5.77 ± 1.04	0.02	0.04
SB 153	10.27 ± 1.72	0.00	0.00
SB 179	8.56 ± 1.26	0.00	0.00
SB 191	5.91 ± 0.69	0.02	0.04

picture in which these massive stars as a population are cospatial with or slightly more dispersed than the gas that they have ionized (e.g. Charlot & Fall 2000; Calzetti 2001; Leitherer et al. 2019), and it is crucially very inconsistent with a scenario in which the smaller *HST/COS* aperture is missing some WR stars captured by the SDSS aperture.

The above experiments indicate that both the UV continuum and the UV stellar wind lines measured by *HST/COS* are consistent with having been produced by the same population of massive stars as captured in the SDSS fibre. We leave a more in-depth exploration of the UV–optical reddening and aperture matching to future work, and proceed with confidence that our comparison of optical nebular emission and UV stellar wind features here is reasonable.

APPENDIX C: CASE-BY-CASE ANALYSIS OF FITS TO THE NEBULAR LINES AND STELLAR WIND FEATURES

Our stellar population synthesis models generally struggle to simultaneously reproduce both the optical nebular line emission and UV stellar wind features detected in these systems. To investigate this in more detail, it is instructive to consider the fits for each galaxy on a case-by-case basis:

SB 9 presents fairly average properties for our sample in the UV, including a C IV profile with absorption depth of -4.4 Å and He II emission at 2.2 Å both near the median for our targets. The gas-phase oxygen abundance inferred from the BEAGLE fits is $12 + \log O/H = 8.39 \pm 0.04$, 0.09 dex higher than inferred from the direct- T_e method above (Table 4).⁹ The model He II profile is in good agreement with that observed before any UV information is incorporated, though the C IV profile is somewhat underestimated in absorption depth and in emission. Adding the stellar wind equivalent widths brings the absorption component of the C IV profile into better agreement

with the data, though the emission is still underestimated. In both fits, the optical WR features are well-modelled by the data, with good agreement to both the He II $\lambda 4686$ and the weak N II and C IV emission. However, in order to match the C IV profile, the BEAGLE models prefer a higher metallicity. The preferred gas-phase O/H is shifted 0.06 dex higher, while the stellar metallicity is shifted even higher, increasing by a factor of 1.4 from $\log_{10}(Z/Z_{\odot}) = -0.37$ to -0.22 . This larger boost to the stellar metallicity relative to the change in the gas-phase is accounted for by an increase in the amount of metals locked-up in dust, with ξ_d nearly doubling from 0.29 to 0.46 , in excess of the solar value $\xi_{d,\odot} = 0.36$ (and near the upper edge of the model range and our prior at 0.5 ; Gutkin et al. 2016). This increase in depletion on to dust allows for an even higher total metallicity Z at fixed O/H. While most of the nebular lines are relatively unaffected by this change, this increase in the preferred model metallicity brings the [O III] $\lambda 4363$ line out of good agreement with the data, lowering its preferred value in the models to 0.8 times that observed.

The stellar wind properties of **SB 49** prove to be even more challenging to reproduce in the models. This system presents the deepest C IV P-Cygni absorption in our sample at -6.51 Å and broad He II emission at 3.20 Å, in-excess of canonical stellar population synthesis predictions (Section 4). The BEAGLE fits to the nebular lines only are in reasonably good agreement with the observed stellar He II profile, while the strength of the C IV P-Cygni absorption and emission is significantly underpredicted. The retrieved gas-phase oxygen abundance for this fit is $12 + \log O/H = 8.53 \pm 0.03$, 0.33 dex higher than found with the direct- T_e method (Tables 4 and 6) though with good agreement found for [O III] $\lambda 4363$. However, incorporating the UV stellar wind equivalent widths in the fits leads to significant discord with these nebular line measurements. In order to match the depth of the observed C IV P-Cygni profile, the models require a significantly higher metallicity; as for SB 9, this is accomplished by a simultaneous increase in both the total metallicity Z and the dust-to-metal mass ratio ξ_d , this time bringing both quantities to the edge of their respective priors (at solar metallicity and $\xi_d \simeq 0.5$). In addition, the model transitions to a preference for an essentially single-age stellar population, with both the start and end times for the current star formation epoch converging to 2 Myr ago. This brings the stellar C IV profile into much better agreement with the data, though the models now slightly overestimate the depth of the absorption trough at intermediate velocities and the strength of emission near systemic (in the latter case, the model predicted emission appears to be in-part nebular, and this disagreement may be partially accounted for by interstellar absorption). This process leaves He II reasonably close to the observed strength (though slightly lowered from the prior estimate due to the general weakening of this feature in the models above $Z/Z_{\odot} \simeq 0.5$; Section 4). However, as for SB 9, this concordance models fail to match the strength of the temperature-sensitive [O III] $\lambda 4363$ line, underpredicting the observed flux by nearly a factor of two. In addition, while the agreement with the optical WR features is very good with the nebular line information alone, the introduction of the UV stellar equivalent width information brings the models into substantial disagreement with the observed profiles. In particular, both the broad He II $\lambda 4686$ and N III $\lambda 4640$ features are substantially overestimated by the models along with broad emission underlying He I $\lambda 5876$; while the observed broad C IV emission at 5808 Å is now undetected in the models. This suggests that the models prefer a solution to the UV stellar wind features relying on very prominent WN stars, which are too high in metallicity or present in too large of numbers to be consistent with the optical data.

⁹However, note that because SB 9 lacks a measurement of [O II] $\lambda 3727$, the BEAGLE fits do not include any O II lines.

The UV spectrum of **SB 61** presents more modest stellar wind features, with both the C IV absorption and He II emission near 2 Å. The nebular-only fit constrains the gas-phase metallicity to be $12 + \log \text{O}/\text{H} = 8.25 \pm 0.05$, only 0.14 dex in-excess of the direct- T_e determination. These optical-only results are already in good agreement with the depth of the C IV P-Cygni absorption, though they slightly underestimate the strength of the stellar emission component and predict a nebular emission component that we do not observe (though again, this comparison is complicated by ISM absorption). In contrast, the stellar He II profile predicted by the models is more prominent than observed. Incorporating the stellar equivalent widths in the fits changes very little, with only a modest increase in the preferred metallicity from $\log_{10} \text{Z}/\text{Z}_{\odot} = -0.57$ to -0.46 and very little change to the predicted C IV profile. Both model predictions show a narrow nebular emission component in the He II profile, of which there is also some indication at low S/N in the observed spectrum. This is the only system in which the model fits predict a clear nebular He II component. This is consistent with the fact that this system has the lowest preferred BEAGLE model metallicity in our sample ($\text{Z}/\text{Z}_{\odot} \simeq 0.3$), in the context of both observed and theoretical evidence towards enhancement in the >54.4 eV photons necessary to power gas recombination emission in this line at lower stellar metallicities (e.g. Senchyna et al. 2019, and references therein). Interestingly, the models incorporating the UV He II equivalent width prefer a stronger broad stellar component and weaker nebular component than observed, as is also evident in the optical He II $\lambda 4686$ profile. This suggests that the models cannot reproduce the strength of the narrow nebular component of this line even at the moderately low metallicity of this system; though evaluating this confidently would require a direct fit to the He II profile rather than just the total equivalent width.

SB 80 (previously studied in S17) reveals a similar picture. With only the nebular line constraints, we constrain the gas-phase oxygen abundance to be $12 + \log \text{O}/\text{H} = 8.33 \pm 0.03$, 0.09 dex higher than the direct- T_e measurement presented in S17. Incorporating the stellar wind strengths in the fit changes little in the derived parameters, with a boost upwards in total metallicity of only 0.08 dex. The fits with the UV stellar equivalent widths also show no significant improvement in the UV, and though C IV is fairly well-fit, He II remains underestimated by the models (though continuum determination near He II in the *HST*/COS data is somewhat uncertain).

The UV spectrum of **SB 119** presents a more significant challenge to reproduce, with prominent stellar He II emission at 4.3 Å. Fitting only the optical lines, we find a gas-phase $12 + \log \text{O}/\text{H} = 8.37 \pm 0.05$ which exceeds the direct- T_e measurement by 0.26 dex. However, this fit significantly underestimates the prominence of both the C IV and He II wind features observed in the *HST*/COS G160M spectrum. Incorporating the stellar wind equivalent widths in the fit brings the models into very good agreement with the full profiles of both features, without shifting the estimated gas-phase oxygen abundance at all. Instead, the current epoch of star formation contracts, with a revised start point of 6.3 Myr rather than 20 and a slightly early truncation 1.3 Myr ago. The inferred stellar metallicity increases by 0.13 dex, but purely through an increase in ξ_d from 0.17 to a significantly supersolar value of 0.46 (near the edge of our prior). The median model-predicted [O III] $\lambda 4363$ flux does decrease somewhat (to 93 per cent of the observed flux), but the posterior distribution remains consistent with the observed value. However, while the UV He II feature is reasonably well-fit, the optical He II $\lambda 4686$ profile is overestimated by the models incorporating the UV stellar equivalent width. Given the relatively low signal-to-noise of the continuum in the UV, this disagreement may be due to an

overestimate of the equivalent width of He II $\lambda 1640$; or this may reflect disagreement in the model prediction of the stellar populations present. In short, our model is able to simultaneously reproduce the set of nebular lines and stellar wind features provided for SB 119, even including its large stellar He II $\lambda 1640$ equivalent width; but only with an unrealistically high ξ_d and with some disagreement in the optical WR features.

SB 125 reveals a C IV P-Cygni absorption equivalent width of -4.6 Å and accompanying stellar He II emission at an intermediate 2.7 Å, both typical for our sample. The nebular line-only fit supports a gas-phase abundance of $12 + \log \text{O}/\text{H} = 8.38 \pm 0.03$, 0.19 dex higher than the direct- T_e result. The observed C IV profile is slightly underestimated by the models, while He II is well-matched. Addition of these stellar equivalent widths to the fit results in increases in both the preferred total metallicity and the corresponding gas-phase oxygen abundance by ~ 0.2 dex, producing another case in which the model [O III] $\lambda 4363$ flux distribution systematically underestimates the observed line flux. In addition, while the agreement with He II $\lambda 1640$ appears unchanged with the addition of the stellar equivalent width measurements to the fit, the optical stellar N III $\lambda 4640$ feature and He II $\lambda 4686$ feature are both overestimated by the models after continuum normalization. This suggests again that the models have resolved the disagreement with the UV stellar features by invoking WN stars at higher metallicity or incidence than supported by the data. Intriguingly, the optical He II profile shows a narrow nebular component in He II which is not visible in the UV He II $\lambda 1640$ profile, even though the contaminating MW Al II $\lambda 1670$ absorption line nearby does not reach the expected line centre for nebular emission in He II $\lambda 1640$. However, we find that given the relatively high reddening measured towards this target from the Balmer decrement ($E(B-V) = 0.25$) and assuming an SMC extinction curve, the non-detection of He II $\lambda 1640$ is consistent with the UV continuum noise and the flux of $\lambda 4686$ given a conservative flux ratio of 10:1 (Hummer & Storey 1987).

SB 126 also presents fairly typical stellar wind properties for our sample, with C IV absorption at -4.0 and 1.9 Å He II emission. The nebular line fits produce a gas-phase oxygen abundance estimate of $12 + \log \text{O}/\text{H} = 8.28 \pm 0.03$, in-excess of the direct- T_e measurement by 0.26 dex. This fit appears in reasonable agreement with the observed broad emission in both the UV and optical He II profiles, but significantly underestimates the strength of the C IV P-Cygni feature. Adding the stellar equivalent widths brings the absorption component of the C IV profile into agreement with the data, but the model still underestimates the strength of the redshifted stellar emission. As before, better matching the UV stellar wind features necessitates a move to higher metallicity, this time by 0.36 dex in Z accompanied by a shift in ξ_d to 0.49 at the edge of the model range. In this case, [O III] $\lambda 4363$ is underestimated by a factor of 0.85 while in addition, the model flux in [O III] $\lambda 5007$ exceeds that observed by a factor of 1.22, leading to an even larger disagreement in the ratio of these lines.

Beginning with **SB 153**, the remaining three targets are in the rarefied class of object with stellar He II equivalent widths comparable to or exceeding 4 Å, and represent a more significant challenge to reconcile. The gas-phase metallicity inferred from the optical-only fit here is $12 + \log \text{O}/\text{H} = 8.36 \pm 0.03$, 0.16 dex larger than estimated with the direct- T_e method. The posterior distribution of the models in the UV is broad with nebular lines only, and somewhat underestimates the strength of the C IV profile. However, adding the stellar wind equivalent widths into the fits brings the models into very good agreement with both C IV and He II. Interestingly, with a relatively modest increase in the inferred metallicity (increasing

from $\log_{10}(Z/Z_{\odot}) = -0.45 \pm 0.05$ to -0.37 ± 0.01) and with a compensating increase in ξ_d to 0.35 ± 0.07 , all of the fitted nebular lines including $\lambda 4363$ are brought into good agreement with the model alongside the stellar wind features. The optical WR features are generally in good agreement with the models, though the observed He II $\lambda 4686$ profile is somewhat lower in equivalent width than preferred by the best-fitting model.

SB 179 presents the largest integrated He II equivalent width in our sample at $4.67 \pm 0.23 \text{ \AA}$. The nebular line fit yields a gas-phase metallicity of $12 + \log O/H = 8.44 \pm 0.03$, only 0.09 dex larger than the direct- T_e determination in **S17**. This nebular line-only fit significantly underestimates the absorption and emission component of the C IV profile and does not match the strength of the observed He II at all. Adding these equivalent width constraints does bring the models into reasonable agreement with the C IV profile. However, the modelled strength of He II remains significantly below that observed, even with a significant reduction of the star-formation period to a brief burst sustained from 4.0 to 1.6 Myr ago. In contrast, while the optical WR features are initially reasonably well-fit, the inclusion of the UV stellar equivalent widths in the fits increases the modelled strength of stellar N III $\lambda 4640$, He II $\lambda 4686$, and C IV $\lambda 5808$ above their observed equivalent widths. In addition, the increase in metallicity to nearly solar ($\log Z/Z_{\odot} = -0.09 \pm 0.1$) again causes [O III] $\lambda 4363$ to fall out of alignment with the data (at a factor of 0.72 times the observed strength).

Finally, **SB 191** reveals a similar challenge to the stellar models, with a total He II equivalent width of $4.02 \pm 0.16 \text{ \AA}$. With optical nebular lines only, we infer a gas-phase oxygen abundance $12 + \log O/H = 8.51 \pm 0.02$, 0.21 dex higher than estimated with the direct- T_e method. The corresponding UV stellar wind line morphology is close to that observed for C IV, but significantly below the very prominent He II emission. We note that SB 179 displays an intriguing He II emission shape, with both an apparent double-peak in the broad emission (also visible in SB 179) and narrow (likely nebular) emission superimposed – we discuss this profile in more detail in Section 5.2.3. Fitting the stellar equivalent widths explicitly brings the models closer, but they remain well below the observed He II emission peak. In contrast, the optical WR features appear reasonably well-fit by the models, though with some indication that He II $\lambda 4686$ is overestimated in the fits including the UV stellar lines. In this case, in contrast to the other systems, we do not observe a significant increase in the preferred model metallicity when the stellar emission is incorporated; the median metallicity increases by only 0.04 dex in this case, likely because it is already near the peak of stellar He II emission in the models at $Z/Z_{\odot} \simeq 0.5$ (Section 4).

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