



**HAL**  
open science

## **A broad-band radio study of PSR J0250+5854: the slowest spinning radio pulsar known**

C.H. Agar, P. Weltevrede, Louis Bondonneau, Jean-Mathias Griessmeier, J.W.T. Hessels, W.J. Huang, A. Karastergiou, M.J. Keith, V.I. Kondratiev, J. Künsemöller, et al.

### ► **To cite this version:**

C.H. Agar, P. Weltevrede, Louis Bondonneau, Jean-Mathias Griessmeier, J.W.T. Hessels, et al.. A broad-band radio study of PSR J0250+5854: the slowest spinning radio pulsar known. Monthly Notices of the Royal Astronomical Society, 2021, 508 (1), pp.1102-1114. 10.1093/mnras/stab2496 . hal-03347488

**HAL Id: hal-03347488**

**<https://hal.science/hal-03347488v1>**

Submitted on 6 Dec 2021

**HAL** is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Copyright

# A broad-band radio study of PSR J0250+5854: the slowest spinning radio pulsar known

C. H. Agar<sup>1</sup>,<sup>\*</sup> P. Weltevrede<sup>1</sup>, L. Bondonneau<sup>2,3</sup>, J.-M. Grießmeier<sup>2,4</sup>, J. W. T. Hessels<sup>5,6</sup>,  
W. J. Huang<sup>7,8,9</sup>, A. Karastergiou<sup>10</sup>, M. J. Keith<sup>1</sup>, V. I. Kondratiev<sup>5,11</sup>, J. Künsemöller<sup>12</sup>, D. Li<sup>13,14,15</sup>,  
B. Peng<sup>13</sup>, C. Sobey<sup>16</sup>, B. W. Stappers<sup>1</sup>, C. M. Tan<sup>17,18</sup>, G. Theureau<sup>2,4,19</sup>, H. G. Wang<sup>7,8,9</sup>,  
C. M. Zhang<sup>13,14</sup>, B. Cecconi<sup>3,4</sup>, J. N. Girard<sup>20</sup>, A. Loh<sup>3</sup> and P. Zarka<sup>4,21</sup>

*Affiliations are listed at the end of the paper*

Accepted 2021 August 29. Received 2021 July 28; in original form 2020 December 1

## ABSTRACT

We present radio observations of the most slowly rotating known radio pulsar PSR J0250+5854. With a 23.5-s period, it is close, or even beyond, the  $P$ - $\dot{P}$  diagram region thought to be occupied by active pulsars. The simultaneous observations with the Five-hundred-metre Aperture Spherical radio Telescope (FAST), the Chilbolton and Effelsberg Low Frequency Array (LOFAR) international stations, and New Extension in Nançay Upgrading loFAR (NenuFAR) represent a five-fold increase in the spectral coverage of this object, with the detections at 1250 (FAST) and 57 MHz (NenuFAR) being the highest and lowest frequency published, respectively, to date. We measure a flux density of  $4 \pm 2 \mu\text{Jy}$  at 1250 MHz and an exceptionally steep spectral index of  $-3.5^{+0.2}_{-1.5}$ , with a turnover below  $\sim 95$  MHz. In conjunction with observations of this pulsar with the Green Bank Telescope and the LOFAR Core, we show that the intrinsic profile width increases drastically towards higher frequencies, contrary to the predictions of conventional radius-to-frequency mapping. We examine polarimetric data from FAST and the LOFAR Core and conclude that its polar cap radio emission is produced at an absolute height of several hundreds of kilometres around 1.5 GHz, similar to other rotation-powered pulsars across the population. Its beam is significantly underfilled at lower frequencies, or it narrows because of the disappearance of conal outriders. Finally, the results for PSR J0250+5854 and other slowly spinning rotation-powered pulsars are contrasted with the radio-detected magnetars. We conclude that magnetars have intrinsically wider radio beams than the slow rotation-powered pulsars, and that consequently the latter's lower beaming fraction is what makes objects such as PSR J0250+5854 so scarce.

**Key words:** polarization – stars: magnetars – stars: neutron – pulsars: individual: PSR J0250+5854.

## 1 INTRODUCTION

Pulsars are rapidly rotating, highly magnetized neutron stars; however, some pulsars rotate significantly less rapidly than others. In this paper, we discuss observations of PSR J0250+5854, a radio pulsar with a period  $P = 23.5$  s discovered by Tan et al. (2018) in the Low Frequency Array (LOFAR) Tied-Array All-Sky Survey (Sanidas et al. 2019). It was detected with the LOFAR high band antenna (HBA) array between 110 and 190 MHz, and with the Green Bank Telescope (GBT) between 300 and 400 MHz. It is the longest period radio pulsar discovered to date, more than twice the period of the second slowest spinning (PSR J2251–3711 at  $P = 12.1$  s; Morello et al. 2020) and almost three times slower than the well-studied 8.5-s pulsar PSR J2144–3933 (Young, Manchester & Johnston 1999). Finding such a slow pulsar is rare, which may be explained by the fact that slower pulsars have lower beaming fractions which make it less likely that they are visible to an observer on Earth, and practical limitations due to selection effects, making

them significantly harder to identify in pulsar survey data if only a small number of pulses are present. The presence of red noise in periodicity searches further hinders their identification (e.g. Lazarus et al. 2015; van Heerden, Karastergiou & Roberts 2017). In addition, slower pulsars lose their ability to create electron–positron pairs and accelerate them sufficiently to produce the detectable coherent radio emission (Sturrock 1971). Therefore, models for radio emission can be constrained by finding the slowest pulsars that still remain active. As discussed in Tan et al. (2018), this means it lies in a relatively empty part of the  $P$ - $\dot{P}$  diagram, beyond the ‘death valley’ as defined by Chen & Ruderman (1993) and the vacuum-gap curvature radiation death line proposed by Zhang, Harding & Muslimov (2000). However, its continued activity is consistent with the partially screened gap model (e.g. Szary 2013). Although PSR J2144–3933 rotates almost three-times faster, its much smaller spin-down rate  $\dot{P} = 4.96 \times 10^{-16} \text{ s s}^{-1}$  makes it more constraining for death line models (Mitra et al. 2020).

In this work, we present the first detection of PSR J0250+5854 at frequencies between 1 and 1.5 GHz using the Five-hundred-metre Aperture Spherical radio Telescope (FAST), along with simultaneous observations with the UK608 and DE601 LOFAR international

\* E-mail: [crispinagar.work@gmail.com](mailto:crispinagar.work@gmail.com)

stations, and the New Extension in Nançay Upgrading loFAR (NenuFAR). The NenuFAR detection at 57 MHz is the lowest frequency detection of PSR J0250+5854 published to date, and the FAST detection is the highest, resulting in an extension by a factor of  $\sim 5$  in spectral coverage of this unique source in the radio domain. The radio frequency evolution is a key part of understanding the pulsar emission mechanism. This is particularly relevant for pulsars such as PSR J0250+5854 which are on the cusp of the death line. Multiwavelength observations can provide information on features including the spectral index (how the flux of the pulsar changes with frequency), and changes in the shape and polarization properties of the radio beam.

Measurements of pulsar radio spectra from a large population began with Sieber (1973), Malofeev & Malov (1980), and Izvekova et al. (1981) using frequencies around and below 100 MHz. Most pulsars were found to have steep spectra which could be modelled with a simple power law  $S_\nu \propto \nu^k$ , where  $S_\nu$  is the mean flux density at some frequency  $\nu$  and  $k$  is the spectral index. For some pulsars deviations from this relation were identified in the form of a turnover at low frequencies which can be attributed to absorption mechanisms, whilst others show a cut-off at high frequencies due to a steepening or break in the spectrum (Sieber 1973). Recent studies of radio pulsar spectral indices show a reasonably broad distribution centred at approximately  $-1.5$  (Bilous et al. 2016; Jankowski et al. 2018).

With its period of 23.5 s, PSR J0250+5854 has an extremely large light cylinder (with a radius  $R_{LC} = cP/2\pi = 1.123 \times 10^6$  km, where  $c$  is the speed of light), and hence a tiny polar cap which connects to the open-field-line region. The diameter of the polar cap is  $D_{PC} \approx 2R\sqrt{R/R_{LC}} \approx 60$  m, where  $R = 10$  km is the canonical neutron star radius (e.g. Sturrock 1971). By comparison, a pulsar with a period  $P = 0.5$  s would have  $D_{PC} \approx 410$  m. This implies that for typical emission heights of hundreds of kilometres, the radio beam of PSR J0250+5854, and hence the duty cycle of its radio pulse, can be expected to be very narrow. Indeed Tan et al. (2018) reported a pulse width of only  $\sim 1^\circ$  at 129, 168, and 350 MHz. The shapes of pulse profiles, even after correcting for propagation effects, are in general observed to be frequency dependent. Often, the profile width decreases with increasing frequency which suggests that higher frequency emission is produced lower in the magnetosphere. This correlation is known as radius-to-frequency mapping (RFM hereafter). RFM was first theorized by Ruderman & Sutherland (1975) – the magnetospheric electron density, hence plasma frequency, is expected to decrease with increasing altitude, thereby predicting that the radio beam expands with decreasing frequency. A number of pulsars have been found to deviate from this relation (e.g. Thorsett 1991; Chen & Wang 2014; Pilia et al. 2016) – this suggests that not necessarily the same magnetic field lines are active at all frequencies (or emission heights), resulting in the appearance and disappearance of profile components with observing frequency (e.g. Cordes 1978; Mitra & Rankin 2002). This can obfuscate the geometrical interpretation of measured profile widths, and is discussed further in Section 4.2. Radio polarization data can help in disentangling these effects, which we further explore for PSR J0250+5854 in Section 3.2, although degeneracies often remain (e.g. Keith et al. 2010).

The structure of this paper is as follows. In Section 2, the new observations are described, followed by a brief explanation of the radio frequency interference (RFI) excision techniques used. The analysis of the data, described in Section 3, is divided into three parts: pulse profile evolution with frequency, polarization properties, and the spectral shape of the pulsar flux density. These results are

then discussed in a broader context in Section 4, and our conclusions are summarized in Section 5.

## 2 OBSERVATIONS

As part of a shared-risk proposal, PSR J0250+5854 was observed on the 2019 May 22 with FAST, a facility built and operated by the National Astronomical Observatories, Chinese Academy of Sciences (Nan et al. 2011; Li et al. 2018). With an effective aperture of 300 m in diameter, it is the world’s largest single-dish radio telescope, and is located in a natural depression in Guizhou Province. The central beam of the high-performance 19-beam receiver operating between 1 and 1.5 GHz (Jiang et al. 2020) was used. Two LOFAR international stations – UK608 (United Kingdom) and DE601 (Germany) – and NenuFAR (France) provided overlapping observations. Chilbolton is home to the UK LOFAR station UK608, formally known as the Rawlings Array and the DE601 station is located at Effelsberg.<sup>1</sup> They each consist of two subarrays: the HBA (110–240 MHz) and low band antenna (LBA; 10–90 MHz; Stappers et al. 2011; van Haarlem et al. 2013), although only the HBAs were used in this project and the bandwidth was limited to 110–190 MHz. NenuFAR at the time of our observations consisted of 52 groups of 19 dual-polarized antennas, operating between 10 and 85 MHz (Zarka et al. 2020). It is located alongside and extends the capabilities of the Nançay LOFAR station (FR606). Table 1 gives a summary of the overlapping observations conducted. The LOFAR international stations were observing over the full duration of the FAST observations, and in the case of NenuFAR significantly longer.

Prior to observing PSR J0250+5854, FAST performed a  $\sim 15$ -min observation of a well-known, bright pulsar PSR J0139+5814 to validate the setup of the observing system. A noise diode signal was injected into the FAST multibeam receiver to facilitate polarization calibration. PSR J0250+5854 was observed for two consecutive hours, interspersed with noise diode observations. Finally, the BL Lacertae object J0303+472 (Véron-Cetty & Véron 2006) was observed for purposes of flux calibration – this source was chosen due to its proximity to PSR J0250+5854.

All data were folded and dedispersed with DSPSR (van Straten & Bailes 2011) using the ephemeris and dispersion measure (DM) reported in Tan et al. (2018) to form a pulse sequence. Dedispersion was done coherently for the DE601, UK608, and NenuFAR data, and incoherently for the FAST data. Flux and polarization calibration were done using PSRCHIVE (Hotan, van Straten & Manchester 2004). Further processing made use of PSRSALSA (Weltevrede 2016)<sup>2</sup> and is described later.

Although not part of the simultaneous observations, this project also made use of an observation of PSR J0250+5854 using the LOFAR Core stations conducted on 2017 October 28. This was part of a run of observations conducted by Tan et al. (2018), but is a different data set to the profiles shown in that paper. This LOFAR Core observation was similarly processed with PSRCHIVE.

### 2.1 Data cleaning techniques

PSR J0250+5854 has a low flux density which, in combination with its extraordinarily long period, makes the analysis susceptible to RFI that affects the baseline level during a rotation period.

<sup>1</sup>DE601 was operated as part of the German LOnG-Wavelength consortium at the time of our observations, in a coherently dedispersed folding mode.

<sup>2</sup><https://github.com/weltevrede/psrsalsa>

**Table 1.** Observation properties of the analysed observations. There is a full overlap of the data for the period during which FAST was recording data for PSR J0250+5854 on the 2019 May 22 [these observations are denoted with an asterisk (\*)]. ‘Resolution’ is the number of samples (pulse longitude ‘bins’) per pulse period. The GBT and LOFAR Core observations were performed on 2017 October 25 and 2017 October 28, respectively.

Observation	Centre freq. (MHz)	Bandwidth (MHz)	No. freq. Channels	Start time (UTC)	No. pulses	Length (hh:mm:ss)	Resolution (No. bins)	Full stokes
FAST*	1250.00	500.0	4096	02:34:07	100	00:39:13	8192	Y
FAST*	1250.00	500.0	4096	03:31:23	153	01:00:01	8192	Y
DE601*	158.55	71.4	488	01:56:19	417	02:43:34	1024	N
UK608*	149.71	95.2	1952	02:02:50	382	02:29:50	8192	N
NenuFAR*	56.54	75.0	384	02:03:18	1528	09:59:21	2048	N
GBT	350.00	100.0	4096	11:54:28	242	01:34:55	8192	N
LOFAR Core	148.93	78.1	400	23:57:16	152	00:59:49	16384	Y

In all data sets, the worst-affected frequency channels were identified and excluded from further analysis. The FAST data were affected by stochastic baseline variations that persisted throughout the observation, with a time-scale somewhat larger than the pulsar’s duty cycle. These baseline variations were removed using PARSALSA by subtracting sinusoids, up to the 23<sup>rd</sup> harmonic of the pulse period, plus a constant offset fitted to the off-pulse region for each rotation of the star in each frequency and Stokes parameter independently. These sinusoids have periods which significantly exceed the duty cycle of the pulsar; hence, the shapes of the pulses were not affected by this process.

The RFI in the DE601 and UK608 observations was very different, appearing as short, bright, impulsive spikes orders of magnitude brighter than the pulsar signal. An effective approach to mitigation was to iteratively clip the brightest samples for each rotation of the star and each frequency channel individually. The clipping is done conservatively to ensure that the pulsar signal is unaffected. With the worst RFI suppressed, the remaining RFI and baseline variations were reduced using the same method described for the FAST data.

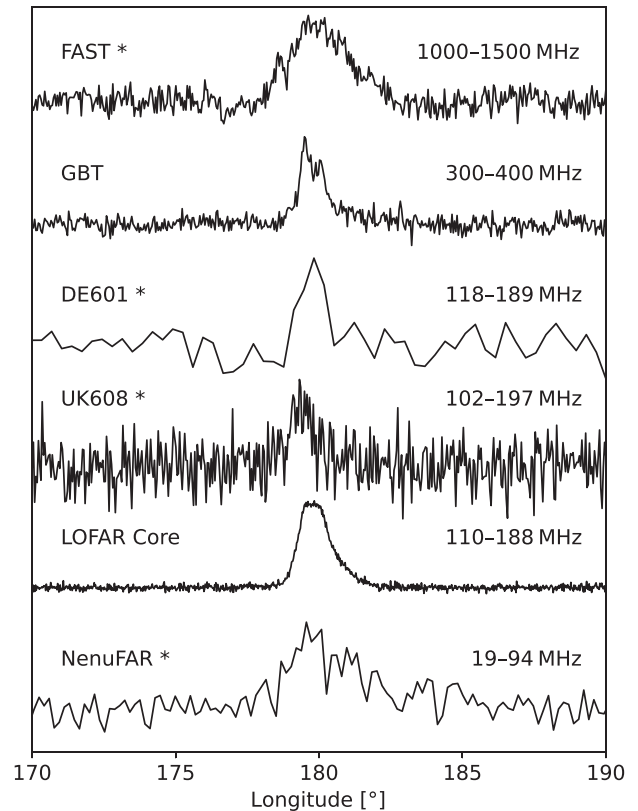
The NenuFAR data were recorded during the commissioning phase of the instrument with a coherent dedispersion pipeline (LUPPI; Bondonneau et al. 2021) operating in single-pulse mode. The observations were folded with DSPSR and a polynomial of degree 2 was subtracted from the baseline of each subintegration to suppress the effect of bandpass variations. The data from two frequency bands were appended after correcting for the appropriate delay.<sup>3</sup> The frequency-resolved observation was cleaned using a modified version of COASTGUARD (Lazarus et al. 2016) and final processing was done with PARSALSA in the same way as with the FAST data.

### 3 ANALYSIS AND RESULTS

#### 3.1 Profile morphology and width evolution

Fig. 1 shows the integrated pulse profile of PSR J0250+5854 as observed by the four telescopes in order of descending frequency. We also include the profile observed by the GBT at 350 MHz from Tan et al. (2018), and a LOFAR Core detection at 149 MHz (an observation from 2017 October 28). The profiles for the LOFAR international stations and NenuFAR are obtained from the full-length observation rather than only the overlap period with the FAST observation to increase the signal-to-noise (S/N) ratio. This is motivated

<sup>3</sup>The observation was recorded in two sub-bands which have a slight offset when dedispersed, and so had to be aligned manually. This is only necessary for older observations conducted before the ‘early science’ phase.

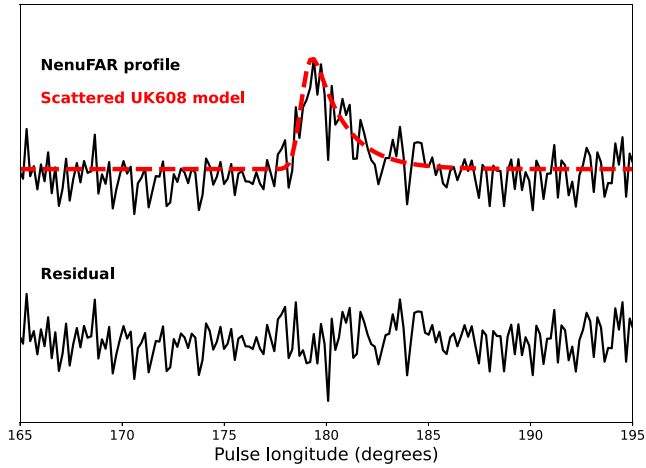


**Figure 1.** The pulse profiles of PSR J0250+5854 at different radio frequencies. The top profile is from FAST (1250 MHz), followed by GBT (350 MHz), then DE601 (154 MHz), UK608 (150 MHz), LOFAR Core (149 MHz), and finally NenuFAR (57 MHz). The FAST, DE601, and UK608 observations are overlapping in time, and are aligned using the known DM and after accounting for geometric delays. The FAST profile is derived from the combination of the two observations. The profiles from the non-simultaneous observations and NenuFAR were visually aligned. The simultaneous observations are denoted with an asterisk (\*).

by the fact that there is no evidence that the profile shapes were changing during these observations. Given the large pulse period, a relatively low number of pulses is recorded, making the profiles less stable (e.g. Helfand, Manchester & Taylor 1975; Rathnasree & Rankin 1995; Liu et al. 2012). This is not a significant issue for the following data analysis given the lack of observed variability.

The profiles of the simultaneous observations in Fig. 1 were aligned by correcting for geometric delays associated with the





**Figure 2.** The NenuFAR profile (top) compared to a smoothed UK608 profile convolved with a scattering tail (dashed red model curve). No significant signal remains in the residuals (bottom).

difference in location of the telescopes (taking right ascension to be  $02^{\text{h}}50^{\text{m}}17^{\text{s}}.78$  and the declination to be  $+58^{\circ}54'01''.3$  as measured by Tan et al. 2018). In addition, the dispersive delay associated with the propagation of the signal through the interstellar medium (ISM) was accounted for by using a DM of  $45.281 \pm 0.003 \text{ cm}^{-3} \text{ pc}$  (Tan et al. 2018).<sup>4</sup> The uncertainty on the DM translates to an uncertainty on the dispersion delay between the highest frequency (FAST; 1250 MHz) and lowest frequency (NenuFAR; 57 MHz) observations of around one pulse longitude bin at the highest resolution shown in Fig. 1 (for the FAST and UK608 data). The NenuFAR data were obtained during commissioning phase and so could not be aligned in this way; hence, the peak of the profile was aligned visually with the UK608 profile peak.

Only the GBT profile has a clear double-peaked profile morphology. Also the LOFAR Core profile with its flat profile peak and the asymmetric FAST profile suggest a more complicated profile structure. It is evident that the profile width at frequencies below that of the FAST observation is significantly narrower, opposite to the expected behaviour by RFM.

The NenuFAR profile, corresponding to the lowest frequency, is broader again and distinctly skewed. Given the steep rise followed by an exponentially decreasing tail, this can be attributed to scattering of the emission in the ISM. This is a strongly frequency-dependent effect with a power-law relationship between the scattering time-scale and frequency, with a power-law index of around  $-4$  (Slee, Dulk & Otrupcek 1980; Lyne & Graham-Smith 2012; Geyer et al. 2017). This suggests that the scattering time-scale for the NenuFAR data is around 50-times greater than at the UK608 centre frequency, which explains why only the NenuFAR profile is significantly affected. The NenuFAR profile is consistent with an intrinsic profile width which is equal to that observed at  $\sim 150$  MHz, albeit broadened by scattering. This is demonstrated in Fig. 2 where the NenuFAR profile is compared with a von Mises function with a width equal to that of the UK608 profile, and convolved with an exponential scattering tail with an e-fold time-scale of 0.1 s. This is consistent with the observed relationship between the DM and scattering time-scale (e.g. Bhat et al. 2004; Ilie, Johnston & Weltevrede 2019). Therefore,

<sup>4</sup>There is no evidence for a change in the DM as the value derived from the NenuFAR data is  $45 \pm 1 \text{ cm}^{-3} \text{ pc}$ , hence consistent with the DM used.

**Table 2.** The profile width at half-maximum ( $W_{50}$ ) of PSR J0250+5854 as a function of frequency, measured by fitting von Mises functions to each profile shown in Fig. 1. These measurements are taken from the profiles as observed, and so include the effect of scatter broadening in the case of the NenuFAR profile.

Telescope	Centre freq. (MHz)	$W_{50}$ ( $^{\circ}$ )
FAST	1250	$2.4 \pm 0.1$
GBT	350	$1.08 \pm 0.07$
DE601	154	$0.9 \pm 0.1$
UK608	150	$0.9 \pm 0.1$
LOFAR Core	149	$1.21 \pm 0.03$
NenuFAR	57	$2.8 \pm 0.4$

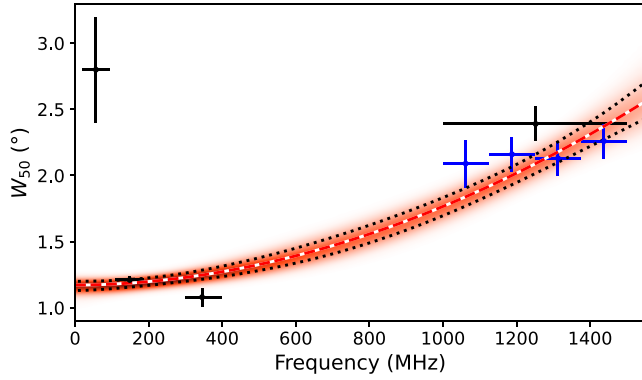
scattering can fully explain the observed frequency evolution of the profile between 60 and 150 MHz (although given that its low S/N ratio makes it nearly impossible to resolve the profile reliably across the frequency band, the possibility of intrinsic profile evolution is not fully excluded). The high S/N LOFAR Core profile also shows a somewhat elongated tail. No evolution of this tail is observed across the bandwidth of the observation, so we conclude that only the NenuFAR profile shows clear evidence for being scatter broadened.

To quantify the pulse broadening at higher frequencies, we measured the profile widths as shown in Fig. 1 by fitting von Mises functions to each profile using PSRSALSA. This smooth mathematical description of the profile allows the width to be measured without being strongly affected by (white) noise. Two components were used to model the higher S/N profiles (LOFAR Core and GBT) but including more than one component for weaker profiles would result in overfitting. The uncertainties on each measurement was calculated using bootstrapping where for each iteration a rotated version of the baseline was added to the profile. This ensures that both the statistical noise arising from the white noise as well as residual baseline variations are accounted for. The estimated full width at half-maximum ( $W_{50}$ ) of the profiles in Fig. 1 are shown in Table 2. To further investigate the frequency evolution of the profile width, the widths were also determined after dividing the FAST data into four frequency sub-bands. Fig. 3 shows the profile width against frequency for the profiles shown in Fig. 1 (black) and the FAST sub-bands (blue).

The evolution of profile width with frequency  $\nu$  was modelled using the relation

$$W_{50} = A\nu^B + C, \quad (1)$$

where  $A$ ,  $B$ , and  $C$  are constants (Thorsett 1991; Chen & Wang 2014). We fitted the function in equation (1) to the profile widths measured from the LOFAR Core, UK608, GBT, and FAST sub-band data. The NenuFAR profile was omitted to avoid scattering in the ISM affecting the results, and the DE601 and UK608 data were omitted because of their low S/N compared to the LOFAR Core data at a similar frequency. The horizontal error bars represent the bandwidth of a given observation. During fitting of equation (1), the frequency of each observation was allowed to vary uniformly within these limits to account for the frequency dependence of the profile width within the observed band. The distribution of fitted trend lines is shown in the red gradient plot in Fig. 3. The black dotted lines bound the 68 per cent confidence interval of the distribution, as a function of frequency. The red dashed line represents the optimal fit to the data, and the power-law exponent of equation (1) is  $1.9 \pm 0.4$ . These findings are discussed further in Section 4.



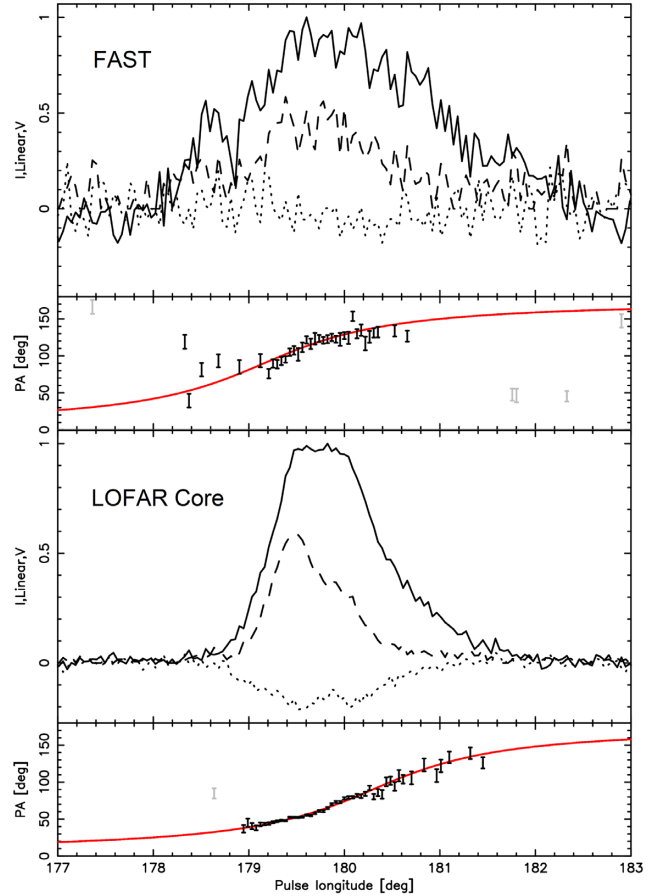
**Figure 3.** Evolution of the profile width of PSR J0250+5854 with observing frequency. Points in black correspond to the profiles shown in Fig. 1 and blue points are the profile widths of the four FAST sub-bands. The red and white dashed line represents the model of frequency evolution in equation (1) which was fitted to the FAST sub-bands, GBT, UK608, and LOFAR Core data (NenuFAR was excluded because it is affected by scattering). A distribution of fits was calculated using bootstrapping techniques, and is represented by the red colour gradient. The black dotted lines indicate the 68 per cent confidence interval of this distribution. The horizontal error bars indicate the bandwidth of a given observation.

### 3.2 Polarization and geometry

The FAST data were calibrated using the pulsed-noise diode signal. After calibration, the polarized FAST pulse profile of PSR J0139+5814 (not shown, see also Section 2) is in excellent agreement with the results of Gould & Lyne (1998) (publicly available on the European Pulsar Network (EPN) data base<sup>5</sup>). The LOFAR Core data were not polarization-calibrated using the LOFAR station beam model, but rather using tied-array addition which incorporates data from different tiles and stations using the station calibration tables to account for the delays between them (more detail can be found in Sobey et al. 2019). The signs of Stokes  $V$  and the position angle (PA) curve had to be flipped in order to agree with convention (e.g. Everett & Weisberg 2001), a correction that was also applied to the FAST data. All data were corrected using the Faraday rotation measure  $RM = -54.65 \pm 0.02 \text{ rad m}^{-2}$ , measured by applying RM synthesis (Brentjens & de Bruyn 2005) to the LOFAR Core polarization data. This is consistent with the RM measured using the FAST data, although this has a larger uncertainty due to the higher observing frequency.

In Fig. 4, the polarized profile of PSR J0250+5854 is shown as observed with FAST (first panel), and the LOFAR Core data (third panel). In both, the solid line is total intensity. The pulse profile has a moderate degree of linear polarization (dashed), which was debiased according to Wardle & Kronberg (1974). There is negative circular polarization (dotted line) in the LOFAR observation, and a hint of the same in the FAST data. The PA ( $\psi$ ) as a function of pulse longitude is shown in the second and fourth panels of Fig. 4, which relates to the Stokes  $Q$ ,  $U$  parameters via  $\psi = 0.5 \arctan(U/Q)$ . Its functional shape can be explained by the Rotating Vector Model (RVM; Radhakrishnan & Cooke 1969), a geometric model which links the observed changes in PA with pulse longitude ( $\phi$ ) to the orientation of the magnetic field lines with respect to the observer.

To fit the RVM, a grid search was conducted over the inclination angle of the magnetic axis,  $\alpha$ , and the impact parameter of the



**Figure 4.** The polarized profile of PSR J0250+5854, observed with FAST (upper plot) and LOFAR Core (lower plot). Total intensity is shown as the solid line, and linear and circular polarization as dashed and dotted line, respectively. The lower panel of each plot shows the variation of PA of linear polarization with pulse longitude. The RVM (red line) was fitted to the LOFAR Core data, and the same curve (with appropriate horizontal and vertical offset applied; see text) is shown for the FAST data. Although more are visible, only the darker PA points were used in fitting the RVM.

observer's line of sight with respect to the magnetic axis,  $\beta$ . For details, see Appendix A and Rookyard, Weltevrede & Johnston (2015b). This was done for the LOFAR Core observation, and the best fit to the observed PA points is shown in Fig. 4 for both the LOFAR Core and FAST data after applying an offset in PA to account for the fact that no absolute PA calibration has been performed, and allowing for a shift of the inflection point in longitude. Only the darker PA points in Fig. 4 were used for fitting, as the PA points in the wings are uncertain and could be affected by orthogonal polarization mode transitions (e.g. McKinnon & Stinebring 2000). The functional shapes of the LOFAR Core and FAST PA data are consistent, as expected when a dipolar field line configuration determines the shape. We therefore will only consider the RVM fit to the higher S/N LOFAR Core data. As will be further discussed in Section 4.3, an offset in the PA inflection point between the FAST and LOFAR Core data  $\Delta\phi_0 = 1.1 \pm 0.8^\circ$  was measured. In Appendix A, a more detailed analysis of the PA data can be found, and the main points are summarized here.

Given the very small duty cycle, limited information is available about how the PA varies with pulse longitude, and as a consequence  $\alpha$  and  $\beta$  are highly correlated. The change of PA with pulse longitude is

<sup>5</sup><http://www.epta.eu.org/epndb/>

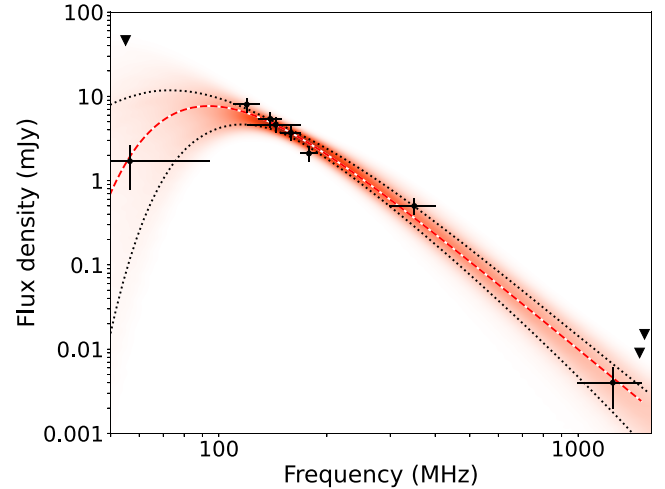
most rapid at the inflection point,  $\sim 55 \text{ deg deg}^{-1}$ , which is predicted by the RVM to be equal to  $\sin \alpha / \sin \beta$  (Komesaroff 1970). This implies that  $\beta$  must be small ( $\lesssim 1.8^\circ$ ) as expected for a detection of a slowly rotating pulsar with a narrow beam directed along the direction of the magnetic axis. The magnetic inclination angle  $\alpha$  is unconstrained from RVM fitting alone.

The measured profile widths provide additional constraints. Assuming the wider profile observed at FAST covers the full extent of the open-field-line region in a dipole geometry and the emission height lies within the range of 200–400 km (e.g. Mitra & Rankin 2002; Johnston & Karastergiou 2019), then this corresponds to a half-opening angle of the beam  $\rho \approx 1 - 2^\circ$  (see Appendix A). This is consistent with what is expected for conal emission (Rankin 1993), see Appendix B. This expectation for  $\rho$  is related to the observed profile width via  $\alpha$  and  $\beta$ , leading to the constraint  $20 \lesssim \alpha \lesssim 50^\circ$ , or  $130 \lesssim \alpha \lesssim 170^\circ$ . The frequency evolution of the beam will be discussed in Section 4.4.

### 3.3 Flux density spectrum

With a factor of  $\sim 5$  increase in spectral coverage with respect to Tan et al. (2018), the radio spectrum of PSR J0250+5854 could be further quantified. Flux calibration of the FAST data was possible by utilizing the observation of the nearby BL Lacertae object J0303+472, which was used as a reference source (see Section 2). This source, also known by the identifier 4C 47.08 (Véron-Cetty & Véron 2006), has a known flux density of 1.8 Jy at a wavelength of 20 cm (approximately 1500 MHz, suitably close to the centre frequency of the FAST data at 1250 MHz) as listed in the VLA Calibrator List.<sup>6</sup> The NASA/IPAC Extragalactic Database (NED)<sup>7</sup> entry for this object contains a list of flux densities of this source at different frequencies from the literature. This reveals a significant scatter in flux density measurements of observations at similar frequencies. To accommodate this, as well as the uncertainty in the intrinsic flux because of interstellar scintillation, we assign an uncertainty of 50 per cent to the flux density. This is consistent with other work on pulsar flux density measurements (e.g. Sieber 1973). The flux density calibration was performed using PSRCHIVE, and we measured a flux density of  $4 \pm 2 \mu\text{Jy}$  for PSR J0250+5854 at 1250 MHz. The S/N ratio of the profile is also consistent with what is predicted for this flux density by the radiometer equation (e.g. Lorimer & Kramer 2005) with known (zenith angle dependent) values for the gain  $G = 14 \text{ K Jy}^{-1}$  and system temperature  $T_{\text{sys}} = 25 \text{ K}$  of FAST<sup>8</sup> (Li et al. 2018). This flux density is below the upper limits at a similar frequency based on non-detections with the Lovell and Nançay telescopes (Tan et al. 2018).

Fig. 5 shows the flux density of PSR J0250+5854 as a function of observing frequency, and includes the flux densities previously measured by Tan et al. (2018). These previous measurements include detections with the GBT, LOFAR HBAs, and a flux density measurement obtained from the LOFAR Two-meter Sky Survey (LoTSS; Shimwell et al. 2017). The detection of PSR J0250+5854 at 57 MHz using NenuFAR marks the lowest frequency detection that is published. The flux density of the pulsar at this frequency was estimated using the radiometer equation, and was found to be  $1.7 \pm 0.9 \text{ mJy}$ , where we have again assigned a 50 per cent



**Figure 5.** The flux density spectrum of PSR J0250+5854, including upper limits (inverted triangles) from previous non-detections (Tan et al. 2018). As in Fig. 3, the horizontal error bars indicate the bandwidth of a given observation. This plot includes the previous flux density measurements from Tan et al. (2018), with the new addition of the NenuFAR and FAST measurements. A power-law relationship with a low-frequency turnover was fitted to the data, and the best fit is indicated by the red and white dashed line. The distribution of acceptable fits is indicated, similar to Fig. 3. The NenuFAR bandwidth extends down to 19 MHz, beyond what is shown in the figure. The fitted parameters from equation (2) are the scaling factor  $b = 0.1_{-0.0}^{+0.3}$ , turnover parameter  $m = 2.1_{-1.2}^{+0.0}$ , critical frequency  $\nu_c = 94 \pm 24 \text{ MHz}$ , and spectral index  $k = -3.5_{-1.4}^{+0.2}$ .

uncertainty. In calibrating these data the elevation of the source and number of antennas in the array were taken into account as they affect the gain, as does the bandpass of the array. The sky background temperature was estimated to be 9050 K at the position of PSR J0250+5854 (which dominates over the receiver temperature of 776 K), found by extrapolating the sky temperature measured at 408 MHz (Haslam et al. 1982) with a spectral index of  $-2.55$  (Lawson et al. 1987; Reich & Reich 1988) to the centre frequency of 56.54 MHz. Full details of the NenuFAR flux calibration procedure are to be published in the instrumentation paper (Zarka et al., in preparation). This measurement indicates that the spectrum of PSR J0250+5854 rolls over at low frequencies, and this explains the upper limit at a similar frequency reported by Tan et al. (2018) based on LOFAR LBA array observations. We further discuss the spectral shape in Section 4.1.

## 4 DISCUSSION

### 4.1 Flux density and spectral index

Tan et al. (2018) were able to detect PSR J0250+5854 over the frequency range between 120 and 350 MHz, and fitted a spectral index of  $-2.6 \pm 0.5$ , which is on the steeper side of the population distribution. However, with neither detection of the pulsar at higher frequencies (1484 and 1532 MHz) using the Lovell and Nançay radio telescopes (respectively), nor a detection with the core LOFAR LBA stations (55 MHz), uncertainty remained over the broad-band shape of the radio spectrum.

PSR J0250+5854 is weak at a centre frequency of 57 MHz as observed with the NenuFAR telescope (compared to its flux density at 150 MHz), which is the result of a spectral turnover (see Section 3.3). This is not unusual in the pulsar population: e.g. Bilous

<sup>6</sup><https://science.nrao.edu/facilities/vla/observing/callist>

<sup>7</sup><https://ned.ipac.caltech.edu/>

<sup>8</sup>See also appendix 2 of [http://english.nao.cas.cn/focus2015/201901/t20190130\\_205104.html](http://english.nao.cas.cn/focus2015/201901/t20190130_205104.html)



et al. (2016) noted that 25 per cent of their low-frequency sample was fitted by a broken power law with a turnover typically around 100 MHz. Furthermore, Jankowski et al. (2018) noted that 21 per cent of their sample deviates from a simple power law, exhibiting mainly broken power laws or low-frequency turnovers. The physical reasons for this are uncertain; however, their analysis suggests that the deviations are partially intrinsic to the pulsar emission or because of magnetospheric absorption processes, and partially due to the environment around the pulsar or the ISM.

To quantify the spectral turnover of PSR J0250+5854, a power law with a low-frequency turnover was fitted to the data using the same model used by Jankowski et al. (2018), which is of the form

$$S_\nu = b \left( \frac{\nu}{\nu_0} \right)^k \exp \left( \frac{k}{m} \left( \frac{\nu}{\nu_c} \right)^{-m} \right), \quad (2)$$

where  $\nu_0 = 500$  MHz is a constant (and arbitrary) reference frequency. The fitted parameters are  $b$ , a constant scaling factor;  $k$ , the spectral index;  $\nu_c$ , the turnover frequency; and  $m$  which determines the smoothness of the transition. The value of  $m$  is expected to be positive, and  $\leq 2.1$ .

With only one flux density measurement below the turnover frequency, the parameters are somewhat ill defined. The optimal fit<sup>9</sup> (Fig. 5, red line) is for an exponent  $m = 2.1$ . This corresponds to the sharpest turnover allowed within the free-free absorption model (see Jankowski et al. 2018 and references therein). The fitted spectral index of  $k = -3.5^{+0.2}_{-1.4}$  is steep compared to the mean found for the pulsar population (Bates, Lorimer & Verbiest 2013 found a mean spectral index of  $-1.4$  with unit standard deviation, whilst Jankowski et al. 2018 found a mean of  $-1.60$  with a standard deviation of 0.54). However, other examples of such steep spectral indices exist, including PSR J1234–6423 which has a broken power-law spectrum with a spectral index of  $-3.8 \pm 0.5$  below  $\sim 1700$  MHz (Jankowski et al. 2018).

Tan et al. (2018) noted that there are occasional bright pulses at 350 MHz in the leading component of the profile of PSR J0250+5854. This behaviour was not seen at around 150 MHz. If equally bright single pulses exist in the FAST frequency band, then they should be comfortably detectable above the level of the thermal noise. However, the residual baseline variations in the single pulse data are such that these bright pulses cannot be confidently detected. It therefore remains to be seen how the erratic nature of the single pulses evolves above frequencies of 350 MHz.

## 4.2 Profile width evolution

As can be seen in Fig. 3, the profile width of PSR J0250+5854 increases with observing frequency, from around  $1^\circ$  at 150 MHz to  $2^\circ$  at 1250 MHz. Equation (1) was fitted to the measured profile widths as a function of frequency, resulting in a power-law index  $B = 1.9 \pm 0.4$  (see Section 3.1). Although it is unusual for pulsars to have a positive index, meaning that their profiles broaden with increasing frequency, there are other examples. Chen & Wang (2014) identified 29 pulsars out of 150 with such a positive index based on profiles from the EPN data base. Given the relatively large uncertainty on both our measured value of  $B$  and those measured by Chen & Wang (2014), the index for PSR J0250+5854 is consistent with 24 out of the 29 pulsars with a reported positive index. None of these 29 pulsars have a

<sup>9</sup>The resulting probability density function of  $m$  is highly clustered at the maximum allowed value of 2.1 which was implemented as a prior. Therefore, no meaningful uncertainty on  $m$  could be assigned.

significantly larger  $B$  than that of PSR J0250+5854. This increase in profile width with frequency is contrary to the expectation from RFM.

Higher frequency radiation can be expected to be produced closer to the neutron star. This follows from models based on curvature radiation from relativistic bunches of particles travelling along the magnetic field lines (e.g. Gil, Lyubarsky & Melikidze 2004; Dyks & Rudak 2015, and references therein), as well as those based on plasma instabilities since both the plasma density and the plasma frequency decrease with increasing altitude (e.g. Hibschan & Arons 2001 and references therein; also Gedalin, Gruman & Melrose 2002). As a consequence, the opening angle of the radio beam can be expected to be narrower at higher frequencies, unless a larger fraction of the open-field-line region becomes active. Pilia et al. (2016) studied 100 pulsars and measured their profile widths at frequencies ranging from tens of megahertz up to 1400 MHz. Only in a few cases was the profile width seen to increase significantly with frequency. In those pulsars, profile broadening was indeed caused by the emergence of new profile components as frequency increased.

Observations at a frequency around 800 MHz could help reveal the reasons for the abnormal frequency evolution of the profile of PSR J0250+5854. The profile morphology of PSR J0250+5854 is indeed complex, with a profile shape skewed in both the LOFAR Core and FAST observations. The GBT profile shows a distinct double-peaked structure with distinct behaviours since Tan et al. (2018) noted that the stronger first component was caused by occasional strong individual pulses. At other frequencies no well-separated profile components are observed. However, the flattened peak in the LOFAR Core profile is suggestive of two blended components of similar intensity. This flattening was not visible for the profiles published in Tan et al. (2018), which is because of the lower S/N. By inspecting all available data, no significant profile shape variability has been detected in LOFAR Core observations of PSR J0250+5854.

## 4.3 Emission height

Constraining the viewing geometry is particularly interesting for slowly rotating pulsars to highlight differences with magnetars (see Section 4.5). As seen in Fig. 4, the inflection point of the PA swing occurs close to the centre of both the FAST and LOFAR Core profiles. The lack of significant relativistic aberration and retardation ( $A/R$ ) shift (Blaskiewicz, Cordes & Wasserman 1991) which moves the inflection point towards, or even beyond, the edge of the profile implies that the emission height must be considerably smaller than the light cylinder radius. This affirms that radio pulsars produce emission at an absolute emission altitude that is relatively constant across the population, rather than being at a constant fraction of  $R_{LC}$ . Given the emission height is not at a constant fraction of  $R_{LC}$  there should be a period dependence of the pulse width (e.g. Rankin 1993). Based on this relationship, Karastergiou & Johnston (2007) proposed that the maximum emission height of radio pulsars at 1.4 GHz is around 1000 km, refined to an absolute height range of 200–400 km irrespective of pulse period (Johnston & Karastergiou 2019; Johnston et al. 2020) – this is the range we assume for the FAST profile (1250 MHz) in Section 3.2.

There is an indication of a difference in the longitude of the inflection point of the PA curve between the LOFAR and FAST frequencies of  $\Delta\phi_0 = 1.1 \pm 0.8^\circ$ . The direction of the shift suggests that lower frequencies are produced higher in the magnetosphere in line with conventional RFM, which would make the narrowness of the low-frequency profiles even more striking. A higher S/N detection at the FAST frequency could make this measurement more significant,



and clarify if there are systematic inconsistencies in the PA swing compared to the prediction of the RVM which could play a role.

#### 4.4 Beam shape evolution and polar cap configuration

The profile width evolution of PSR J0250+5854 with frequency is strong, and opposite compared to what is expected from conventional RFM. As pointed out in Section 4.3, such behaviour is often linked to the emergence of new profile components as frequency increases. Although for PSR J0250+5854 there is no strong evidence of the emergence of a new profile component, it is clear the profile morphology is complex (Section 3.1). Under the assumption that emission height increases with decreasing frequency, the implication is that at lower frequencies the beam is underfilled compared to the open-field-line region. This means that only a small fraction of the field lines are significantly active and emission is more concentrated to a specific region of the polar cap. This could be associated with a frequency dependence in the active patches in a patchy beam model (e.g. Lyne & Manchester 1988; Karastergiou & Johnston 2007). Here, we outline other alternative models of the beam structure that could explain our observations.

In the framework of the core–cone model (e.g. Rankin 1983a,b; Radhakrishnan & Rankin 1990; Rankin 1993), the broadening of the pulse profile with increasing frequency can be explained by so-called conal outriders. Here, the narrower central core component dominates at low frequencies. However, the wider conal components with a shallower spectrum become more prominent at higher frequencies. The rise of conal outriders at FAST frequencies then must be blended with the core component to explain the wide single-peaked profile. In this scenario, it is expected that the profile would evolve into a wide, double-peaked profile at even higher radio frequencies. The bifurcated GBT profile in Fig. 1 would therefore not be a consequence of conal emission, but should instead be associated with a magnetospheric absorption feature of a core–single profile, which are known to occur between 200 and 800 MHz (e.g. Rankin 1983b, 1986). In Appendix B, the viewing geometry is derived in this framework, leading to results consistent with Section 3.2 and Appendix A. Unlike most pulsars with conal outriders (e.g. Rankin, Stinebring & Weisberg 1989), PSR J0250+5854 has a low energy-loss rate,  $\dot{E} \sim 10^{29} \text{ erg s}^{-1}$  (compared to  $\dot{E} \gtrsim 10^{32} \text{ erg s}^{-1}$ ).

A non-uniform emission height could amplify the effects of a frequency dependence in emissivity across the open-field-line region. There is evidence for the emission height to be larger near the rim of the open-field-line region (e.g. Gupta & Gangadhara 2003; Weltevrede & Johnston 2008, and references therein). More recently, Rankin, Olszanski & Wright (2020) linked the core–cone model to the plasma pair multiplicity model of Timokhin & Harding (2015) in which the pair production front can form a cup-like structure in the open-field-line region. The increase in emission height towards the rim of the open-field-line region further increases the beaming fraction of emission. With the suggestion that the emission height for PSR J0250+5854 at LOFAR frequencies is higher compared to FAST frequencies (see Section 4.3) this may not be a significant effect.

Chen & Wang (2014) conclude that in a number of pulsars the widening of the profile at higher frequencies could not be ascribed to structures consistent with the core–cone model. This also appears to be the case in a small subgroup of pulsars studied by Pilia et al. (2016). Like PSR J0250+5854, these pulsars do not show well-separated profile peaks at the highest frequencies. The expectation from conventional RFM is based on emission being produced over a wide range of altitudes, with each height producing narrowband emission.

On the other hand, if a narrow range of emission heights generates broad-band emission, the observed spectrum will be different. This led Chen & Wang (2014) to suggest that fan beams could accommodate anti-RFM-like behaviour. Broad-band emission is incorporated into the fan beam model (Michel 1987; Dyks, Rudak & Demorest 2010; Dyks & Rudak 2012, 2013; Wang et al. 2014) where emission is produced along magnetic flux tubes that extend out from the pole in a fan-like structure. A single fan, hence spark, would be required to explain the beam structure of PSR J0250+5854. Support for this beam structure is found in observations of the precessing pulsars J1141–6545 and J1906+0746 (Manchester et al. 2010; Desvignes et al. 2013). Following the suggestion of Michel (1987) that each flux tube may have its own spectrum, Chen et al. (2007) argued that the emission spectrum may not be homogeneous across a flux tube. In particular, they argue that pulsars which show pulse broadening with increasing frequency may have a flattening emission spectrum away from the magnetic axis, as supported by their simulations. Broad-band emission in flux tubes with a location-varying spectral index follows naturally from particle-in-cell simulations of vacuum-gap pair-production (Timokhin 2010) which predict that the momentum spectrum of the secondary plasma is not necessarily monotonic as a function of height within the magnetosphere, and so a given observed frequency cannot be assigned to a unique altitude.

As highlighted in the introduction, the slow rotation of PSR J0250+5854 implies a tiny polar cap connected to the open-field-line region. One can wonder if there is enough space to fully develop the type of beam complexity encountered in more typical pulsars. In the model of Mitra et al. (2020), it is predicted that the polar cap of PSR J2144–3933 is only large enough to support a single pair-production site – known as a ‘spark’ – hence its profile would be a single component. For PSR J0250+5854 (which although slower has a much larger  $\dot{P}$ ), their model predicts that the footprint of a spark is one-fifth of the area of the polar cap, meaning that up to three sparks could be supported if they are packed tightly. Moreover, sparks are also required to be separated from one another due to screening effects (e.g. Gil & Sendyk 2000), meaning that the polar cap may only be able to support a single spark. This single spark as proposed for PSR J2144–3933, may not circulate about the magnetic axis in a way theorized for multispark systems (Ruderman & Sutherland 1975). Such a single-spark scenario may explain why while the profile width evolves as a function of frequency, the profile essentially remains single peaked.

#### 4.5 Comparison to other slow pulsars and magnetars

The extremely long period of PSR J0250+5854 places it on the far right-hand side of the  $P$ – $\dot{P}$  diagram, in an area largely inhabited by magnetars and X-ray dim isolated neutron stars (XDINSSs). These objects are detected only as soft thermal X-ray sources without radio counterparts. Of the seven brightest XDINSSs, five have high magnetic dipole fields of the order of  $10^{13}$ – $10^{14}$  G which may mean they are related to magnetars (Haberl 2007; van Kerkwijk & Kaplan 2007). Despite the fact that magnetars form a distinct class of objects with much greater spin-down rates, and hence much higher rates of loss of rotational energy  $\dot{E}$ , they may evolve with time towards the parameter space occupied by the slow pulsars (e.g. Viganò et al. 2013). However, to date PSR J0250+5854 remains undetected in X-rays, despite a dedicated *Swift* X-ray telescope observation, which makes it difficult to confirm a connection between it and XDINSS (see Tan et al. 2018 for details). Similarly, PSR J0250+5854 has not shown any magnetar-like behaviour such as bursts, or large radio variability

**Table 3.** Parameters (period  $P$ , spin-down rate  $\dot{P}$ ) of the three slowest spinning radio pulsars (top three sources) and the five magnetars (lower five sources) known to produce pulsed radio emission. The profile width of the radio pulsars are measured values of  $W_{50}$  taken from this work and the referenced literature. The widths of the magnetar profiles are estimated based on the full extent over which significant emission was visible in the published profiles (at the reference frequency) in order to capture the complexity of the magnetar profiles.

Object	$P$ (s)	$\dot{P}$ ( $\text{s s}^{-1}$ )	Profile width ( $^\circ$ )	Ref. freq. (MHz)	References
PSR J0250+5854	23.5	$2.72 \times 10^{-14}$	2.4	1250	This work
PSR J2144–3933	8.5	$4.96 \times 10^{-16}$	0.8	1400	1, 2
PSR J2251–3711	12.1	$1.31 \times 10^{-14}$	1.2	1382	3
1E 1547.0–5408	2.1	$2.32 \times 10^{-11}$	90	6600	4, 5
PSR J1622–4950	4.3	$1.70 \times 10^{-11}$	190	1400	6, 7
PSR J1745–2900	3.8	$6.80 \times 10^{-12}$	15	2400	8
XTE J1810–197	5.5	$1.02 \times 10^{-11}$	35	1400	9, 10, 11, 12
Swift J1818.0–1607	1.4	$9 \times 10^{-11}$	20	1548	13, 14, 15

*Note.* References: (1) Young et al. (1999); (2) Mitra et al. (2020); (3) Morello et al. (2020); (4) Camilo et al. (2007c); (5) Camilo et al. (2008); (6) Levin et al. (2010); (7) Levin et al. (2012); (8) Eatough et al. (2013); (9) Camilo et al. (2006); (10) Camilo et al. (2007a); (11) Kramer et al. (2007); (12) Levin et al. (2019); (13) Esposito et al. (2020); (14) Lower et al. (2020); (15) Champion et al. (2020).

as of yet. Furthermore, PSR J0250+5854 is located squarely in the Galactic plane which argues that it is still relatively young.

Nevertheless, since the pulse period is a key factor controlling the width of radio pulse profiles, it is worth comparing PSR J0250+5854 with the radio-emitting magnetars and other slowly spinning rotation-powered radio pulsars. This comparison can highlight what other parameters play a role in the radio beam geometry of these slowly rotating objects. Aside from PSR J0250+5854, the two other slowest spinning known radio pulsars are PSRs J2251–3711 ( $P = 12.1$  s) and J2144–3933 ( $P = 8.5$  s). There are five known magnetars for which pulsed-radio emission has been detected. A summary of their properties is shown in Table 3. Although the focus here will be the differences in profile widths, there are other differences such as the spectra of magnetars being radically different (Camilo et al. 2007a,c; Kramer et al. 2007; Keith et al. 2011; Champion et al. 2020; Lower et al. 2020), and their radio emission being much more transient with periods of activity and strongly changing profile shapes (Serylak et al. 2009; Dai et al. 2018, 2019; Levin et al. 2019; Lower et al. 2021).

The three slowest spinning radio pulsars have long periods and narrow, fairly simple pulse profiles. Measurements of  $W_{50}$  are published for these profiles and are representative of the overall profile width. Despite being the slowest spinning of the three, PSR J0250+5854 has the widest profile by a factor of around two. If only the period determines the width, the inverse would be expected. This suggests that besides PSR J0250+5854 (see Section 4.3) underfilling of the beam may also play a role in the other slow pulsars. In contrast, the magnetars have more complex profiles with distinct components, which means that  $W_{50}$  is not always representative of the overall profile width. Therefore, for the magnetars in Table 3, the full span (rounded to the nearest five degrees) over which emission is seen in the published profiles is reported.

There is a stark contrast between the profile widths of the magnetars compared to the slow pulsars, much more than can be expected from just the differences in  $P$  (and hence  $R_{\text{LC}}$ ). There are three potential geometric explanations: (1) for the slowly rotating pulsars only a tiny fraction of the open-field-line region is active, (2) all magnetars have a magnetic axis almost aligned with the rotation axis, and (3) all magnetars have atypically wide beams due to large emission heights or otherwise. We will argue that only options (2) and (3) are viable, and that option (4) plays a more significant role. We will not consider the very unlikely coincidence that all three

slowly rotating pulsars are observed with extremely grazing lines of sight with respect to the radio beam.

The active fraction of the open-field-line region needs to be very small for the slow pulsars if it is to be the main reason why the slow pulsars have such narrow beams compared to the magnetars. This seems unlikely, as it would not explain why there are no slow pulsars with multiple narrow profile components spread over a similar fraction of the rotation period for which magnetars show emission. The complex magnetar profiles often exhibit individual components which are much wider than the full profiles of the slow pulsars.

If magnetars have very aligned radio beams with respect to their rotation axis (small  $\alpha$ ), the observer’s line of sight would spend a larger fraction of the time within the beam. The process of Section 3.2 (detailed in Appendix A) can be used to estimate how extreme the alignment of the magnetars should be in order to explain their wide pulse profiles. Taking the mean magnetar period and median profile width from Table 3 with an emission height of 400 km, such an object requires  $\alpha \lesssim 11^\circ$  to produce profiles of the observed width. However, there is little evidence for this. Polarization studies have given a variety of  $\alpha$  value for the magnetars ranging from near aligned to almost orthogonal (Camilo et al. 2007b, 2008; Kramer et al. 2007; Levin et al. 2012; Lower et al. 2021). In addition, highly aligned magnetars are difficult to reconcile with the large modulation of the thermal X-rays, as was highlighted for XTE J1810–197 (Gotthelf & Halpern 2007; Perna & Gotthelf 2008) and 1E 1547.0–5408 (Israel et al. 2010). Furthermore, if magnetars evolve into slowly spinning rotation-powered pulsars, no difference in their  $\alpha$  distribution can be expected. For PSR J0250+5854, for example, we have shown that it is unlikely that  $\alpha$  is very small. Therefore, it is concluded that the emission height must play a significant role in explaining the magnetar radio profile widths.

If the emission heights are the dominant reason for the magnetars having wider radio profiles, they need to be  $\sim 20$ -times larger (around 10 000 km) compared to the slowly spinning rotation-powered pulsars. Alternatively, if their radio beams are confined by last open field lines which close within the light cylinder (as suggested by detailed simulations by Spitkovsky 2006; see also the discussion of ‘Y-points’ by Craig 2014), or if magnetic field sweepback plays a large role (e.g. Craig & Romani 2012) then their polar beams will be wider as well. Therefore, one cannot distinguish between large emission heights and extended open-field-line regions

(e.g. Rookyard, Weltevrede & Johnston 2015a; Rookyard et al. 2015b). In such a scenario, the open-field-line regions of magnetars would need to be  $\sim 5$ -times larger than predicted for a static dipole field. In either case, it would imply that slowly spinning rotation-powered pulsars have dramatically reduced beaming fractions compared to magnetars which therefore play an important role in the deficit of observed slow pulsars near the death valley.

The large implied difference in the beaming fraction between the two classes of objects is unexpected, given the weak  $\dot{P}$  dependence of the pulse widths observed for the normal pulsar population (e.g. Kijak & Gil 2003; Johnston & Karastergiou 2019). This implies that for these slowly rotating objects, the role of  $\dot{P}$  in governing the beaming fraction is much larger than for the normal pulsar population. This could potentially be facilitated by the incredible strengths of the magnetar magnetic fields.

## 5 CONCLUSIONS

We have obtained the highest and lowest frequency radio detections of PSR J0250+5854, the most slowly rotating radio-emitting pulsar known, using simultaneous observations from 57 to 1250 MHz. The highest frequency detection with FAST (1250 MHz) shows that the spectrum is exceptionally steep with a spectral index of  $-3.5^{+0.2}_{-1.4}$  and the lowest frequency detection with NenuFAR (57 MHz) reveals a spectral turnover below 95 MHz. While PSR J0250+5854 is slow, it is relatively close to the ‘death valley’ and the fact it is still active in radio can be accommodated in current emission models. The pulse profile shows narrowing at lower frequencies, contrary to the expectations of RFM. This decrease in beaming fraction is suggestive of a reduction in the filling fraction of the beam, or disappearing conal outriders. The polarization information of LOFAR Core data at 150 MHz and FAST data at 1250 MHz was used to show that the line-of-sight impact parameter  $\beta$  is very small, passing within  $1.8^\circ$  of the magnetic axis. This confirms that the radio beam is very narrow, as expected for such a slow pulsar. Furthermore, the lack of a delay between the profile peak and PA curve inflection point implies that the emission height of PSR J0250+5854 at 1250 MHz is low, consistent with those found for other non-recycled pulsars. Finally, we draw comparisons between other slow pulsars, PSR J0250+5854, and the five known magnetars with pulsed-radio emission which have the most similar pulse periods in the known pulsar population. To explain the significantly broader magnetar profiles, magnetic alignment may play a role; however, the main reason is likely to be either considerably expanded open-field-line regions or substantially larger emission heights for magnetars.

## ACKNOWLEDGEMENTS

This research has made use of the NED which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. Part of this research has made use of the European Pulsar Network Database of Pulsar Profiles maintained by the University of Manchester. Pulsar research at Jodrell Bank Centre for Astrophysics and Jodrell Bank Observatory is supported by a consolidated grant from the UK Science and Technology Facilities Council (STFC). This research is also supported by the National Natural Science Foundation of China (NSFC; grant no. 11988101, U1938117, U1731238, and 11703003).

This paper is based on data from the German LOng-Wavelength (GLOW) array, which is part of the International LOFAR Telescope (ILT) which is designed and built by Netherlands Institute for Radio

Astronomy (ASTRON; van Haarlem et al. 2013). Specifically, we used the Effelsberg (DE601) station funded by the Max-Planck-Gesellschaft. The observations of the German LOFAR stations were carried out in the stand-alone GLOW mode which is technically operated and supported by the Max-Planck-Institut für Radioastronomie, the Forschungszentrum Jülich, Bielefeld University, by the Federal Ministry of Education and Research (BMBF) Verbundforschung project D-LOFAR III (grant number 05A14PBA) and by the states of Nordrhein-Westfalen and Hamburg.

This paper is based on data obtained using the NenuFAR radio telescope. The development of NenuFAR has been supported by personnel and funding from: Station de Radioastronomie de Nançay, Centre national de la recherche scientifique – Institut national des sciences de l’Univers (CNRS-INSU), Observatoire de Paris-PSL, Université d’Orléans, Observatoire des Sciences de l’Univers en région Centre, Région Centre-Val de Loire, Domaine d’Intérêt Majeur ‘Astrophysique et conditions d’apparition de la vie’ (DIM-ACAV) and DIM-ACAV+ of Région Ile de France, and Agence Nationale de la Recherche. We acknowledge the use of the Nançay Data Center computing facility (Centre de Données de Nançay; CDN). The CDN is hosted by the Station de Radioastronomie de Nançay in partnership with Observatoire de Paris, Université d’Orléans, L’Observatoire des Sciences de l’Univers en région Centre – Val de Loire (OSUC), and the French Centre National de la Recherche Scientifique (CNRS). The CDN is supported by the Region Centre Val de Loire, département du Cher. The Nançay Radio Observatory is operated by the Paris Observatory, associated with the CNRS.

This paper is based on data obtained with the ILT under project code DDT8\_004. LOFAR (van Haarlem et al. 2013) is the Low Frequency Array designed and constructed by ASTRON. It has observing, data processing, and data storage facilities in several countries, that are owned by various parties (each with their own funding sources), and that are collectively operated by the ILT foundation under a joint scientific policy. The ILT resources have benefited from the following recent major funding sources: CNRS-INSU, Observatoire de Paris, and Université d’Orléans, France; BMBF, the Ministry of Innovation, Science and Research of the German State of North Rhine-Westphalia (MIWF-NRW), and Max-Planck-Gesellschaft zur Förderung der Wissenschaften e. V. (MPG), Germany; Science Foundation Ireland (SFI) and Department of Business, Enterprise and Innovation (DBEI), Ireland; Dutch Research Council (Nederlandse Organisatie voor Wetenschappelijk Onderzoek; NWO), the Netherlands; and The Science and Technology Facilities Council, UK.

JWTH acknowledges funding from an NWO Vici grant (‘AstroFlash’).

## DATA AVAILABILITY

The data underlying this article will be shared on reasonable request to the corresponding author.

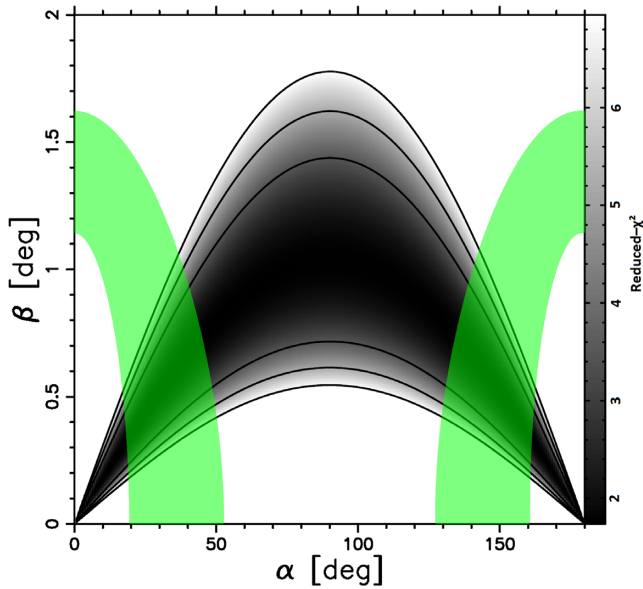
## REFERENCES

- Bates S. D., Lorimer D. R., Verbiest J. P. W., 2013, *MNRAS*, 431, 1352
- Bhat N. D. R., Cordes J. M., Camilo F., Nice D. J., Lorimer D. R., 2004, *ApJ*, 605, 759
- Bilou A. V. et al., 2016, *A&A*, 591, A134
- Blaskiewicz M., Cordes J. M., Wasserman I., 1991, *ApJ*, 370, 643
- Bondonneau L. et al., 2021, *A&A*, 652, A34
- Brentjens M. A., de Bruyn A. G., 2005, *A&A*, 441, 1217



- Camilo F., Ransom S. M., Halpern J. P., Reynolds J., Helfand D. J., Zimmerman N., Sarkissian J., 2006, *Nature*, 442, 892
- Camilo F., Reynolds J., Johnston S., Halpern J. P., Ransom S. M., van Straten W., 2007a, *ApJ*, 659, L37
- Camilo F. et al., 2007b, *ApJ*, 663, 497
- Camilo F., Ransom S. M., Halpern J. P., Reynolds J., 2007c, *ApJ*, 666, L93
- Camilo F., Reynolds J., Johnston S., Halpern J. P., Ransom S. M., 2008, *ApJ*, 679, 681
- Champion D. et al., 2020, *MNRAS*, 498, 6044
- Chen K., Ruderman M., 1993, *ApJ*, 402, 264
- Chen J. L., Wang H. G., 2014, *ApJS*, 215, 11
- Chen J.-L., Wang H.-G., Chen W.-H., Zhang H., Liu Y., 2007, *Chin. J. Astron. Astrophys.*, 7, 789
- Cordes J. M., 1978, *ApJ*, 222, 1006
- Craig H. A., 2014, *ApJ*, 790, 102
- Craig H. A., Romani R. W., 2012, *ApJ*, 755, 137
- Dai S. et al., 2018, *MNRAS*, 480, 3584
- Dai S. et al., 2019, *ApJ*, 874, L14
- Desvignes G., Kramer M., Cognard I., Kasian L., van Leeuwen J., Stairs I., Theureau G., 2013, in van Leeuwen J., ed., Proc. IAU Symp. 291, Neutron Stars and Pulsars: Challenges and Opportunities After 80 Years. Cambridge Univ. Press, Cambridge, p. 201
- Dyks J., Rudak B., 2012, *MNRAS*, 420, 3403
- Dyks J., Rudak B., 2013, *MNRAS*, 434, 3061
- Dyks J., Rudak B., 2015, *MNRAS*, 446, 2505
- Dyks J., Rudak B., Demorest P., 2010, *MNRAS*, 401, 1781
- Eatough R. P. et al., 2013, *Nature*, 501, 391
- Esposito P. et al., 2020, *ApJ*, 896, L30
- Everett J. E., Weisberg J. M., 2001, *ApJ*, 553, 341
- Gedalin M., Gruman E., Melrose D. B., 2002, *MNRAS*, 337, 422
- Geyer M. et al., 2017, *MNRAS*, 470, 2659
- Gil J. A., Sendyk M., 2000, *ApJ*, 541, 351
- Gil J., Gronkowski P., Rudnicki W., 1984, *A&A*, 132, 312
- Gil J., Lyubarsky Y., Melikidze G. I., 2004, *ApJ*, 600, 872
- Gotthelf E. V., Halpern J. P., 2007, *Ap&SS*, 308, 79
- Gould D. M., Lyne A. G., 1998, *MNRAS*, 301, 235
- Gupta Y., Gangadhara R. T., 2003, *ApJ*, 584, 418
- Haberl F., 2007, *Ap&SS*, 308, 181
- Haslam C. G. T., Salter C. J., Stoffel H., Wilson W. E., 1982, *A&AS*, 47, 1
- Helfand D. J., Manchester R. N., Taylor J. H., 1975, *ApJ*, 198, 661
- Hibschman J. A., Arons J., 2001, *ApJ*, 560, 871
- Hotan A. W., van Straten W., Manchester R. N., 2004, *Publ. Astron. Soc. Aust.*, 21, 302
- Ilie C. D., Johnston S., Weltevrede P., 2019, *MNRAS*, 483, 2778
- Israel G. L. et al., 2010, *MNRAS*, 408, 1387
- Izvekova V. A., Kuzmin A. D., Malofeev V. M., Shitov I. P., 1981, *Ap&SS*, 78, 45
- Jankowski F., van Straten W., Keane E. F., Bailes M., Barr E. D., Johnston S., Kerr M., 2018, *MNRAS*, 473, 4436
- Jiang P. et al., 2020, *Res. Astron. Astrophys.*, 20, 064
- Johnston S., Karastergiou A., 2019, *MNRAS*, 485, 640
- Johnston S., Smith D. A., Karastergiou A., Kramer M., 2020, *MNRAS*, 497, 1957
- Karastergiou A., Johnston S., 2007, *MNRAS*, 380, 1678
- Keith M. J., Johnston S., Weltevrede P., Kramer M., 2010, *MNRAS*, 402, 745
- Keith M. J., Johnston S., Levin L., Bailes M., 2011, *MNRAS*, 416, 346
- Kijak J., Gil J., 2003, *A&A*, 397, 969
- Komesaroff M. M., 1970, *Nature*, 225, 612
- Kramer M., Stappers B. W., Jessner A., Lyne A. G., Jordan C. A., 2007, *MNRAS*, 377, 107
- Lawson K. D., Mayer C. J., Osborne J. L., Parkinson M. L., 1987, *MNRAS*, 225, 307
- Lazarus P. et al., 2015, *ApJ*, 812, 81
- Lazarus P., Karuppusamy R., Graikou E., Caballero R. N., Champion D. J., Lee K. J., Verbiest J. P. W., Kramer M., 2016, *MNRAS*, 458, 868
- Levin L. et al., 2010, *ApJ*, 721, L33
- Levin L. et al., 2012, *MNRAS*, 422, 2489
- Levin L. et al., 2019, *MNRAS*, 488, 5251
- Li D. et al., 2018, *Proc. IEEE*, 19, 112
- Liu K., Keane E. F., Lee K. J., Kramer M., Cordes J. M., Purver M. B., 2012, *MNRAS*, 420, 361
- Lorimer D. R., Kramer M., 2005, Handbook of Pulsar Astronomy. Cambridge Univ. Press, Cambridge
- Lower M. E., Shannon R. M., Johnston S., Bailes M., 2020, *ApJ*, 896, L37
- Lower M. E., Johnston S., Shannon R. M., Bailes M., Camilo F., 2021, *MNRAS*, 502, 127
- Lyne A., Graham-Smith F., 2012, Pulsar Astronomy, 4th edn. Cambridge Univ. Press, Cambridge
- Lyne A. G., Manchester R. N., 1988, *MNRAS*, 234, 477
- McKinnon M. M., Stinebring D. R., 2000, *ApJ*, 529, 435
- Malofeev V. M., Malov I. F., 1980, *SvA*, 24, 54
- Manchester R. N. et al., 2010, *ApJ*, 710, 1694
- Michel F. C., 1987, *ApJ*, 322, 822
- Mitra D., Rankin J. M., 2002, *ApJ*, 577, 322
- Mitra D., Basu R., Melikidze G. I., Arjunwadkar M., 2020, *MNRAS*, 492, 2468
- Morello V. et al., 2020, *MNRAS*, 493, 1165
- Nan R. et al., 2011, *Int. J. Mod. Phys. D*, 20, 989
- Perna R., Gotthelf E. V., 2008, *ApJ*, 681, 522
- Pilia M. et al., 2016, *A&A*, 586, A92
- Radhakrishnan V., Cooke D. J., 1969, *Astrophys. Lett.*, 3, 225
- Radhakrishnan V., Rankin J. M., 1990, *ApJ*, 352, 258
- Rankin J. M., 1983a, *ApJ*, 274, 333
- Rankin J. M., 1983b, *ApJ*, 274, 359
- Rankin J. M., 1986, *ApJ*, 301, 901
- Rankin J. M., 1990, *ApJ*, 352, 247
- Rankin J. M., 1993, *ApJ*, 405, 285
- Rankin J. M., Stinebring D. R., Weisberg J. M., 1989, *ApJ*, 346, 869
- Rankin J. M., Olszanski T. E. E., Wright G. A. E., 2020, *ApJ*, 890, 151
- Rathnasree N., Rankin J. M., 1995, *ApJ*, 452, 814
- Reich P., Reich W., 1988, *A&A*, 196, 211
- Rookyard S. C., Weltevrede P., Johnston S., 2015a, *MNRAS*, 446, 3356
- Rookyard S. C., Weltevrede P., Johnston S., 2015b, *MNRAS*, 446, 3367
- Ruderman M. A., Sutherland P. G., 1975, *ApJ*, 196, 51
- Sanidas S. et al., 2019, *A&A*, 626, A104
- Serylak M. et al., 2009, *MNRAS*, 394, 295
- Shimwell T. W. et al., 2017, *A&A*, 598, A104
- Sieber W., 1973, *A&A*, 28, 237
- Slee O. B., Dulk G. A., Otrupcek R. E., 1980, *Proc. Astron. Soc. Aust.*, 4, 100
- Sobey C. et al., 2019, *MNRAS*, 484, 3646
- Spitkovsky A., 2006, *ApJ*, 648, L51
- Stappers B. W. et al., 2011, *A&A*, 530, A80
- Sturrock P. A., 1971, *ApJ*, 164, 529
- Szary A., 2013, PhD thesis, Univ. Zielona Góra
- Tan C. M. et al., 2018, *ApJ*, 866, 54
- Thorsett S. E., 1991, *ApJ*, 377, 263
- Timokhin A. N., 2010, *MNRAS*, 408, 2092
- Timokhin A. N., Harding A. K., 2015, *ApJ*, 810, 144
- van Haarlem M. P. et al., 2013, *A&A*, 556, A2
- van Heerden E., Karastergiou A., Roberts S. J., 2017, *MNRAS*, 467, 1661
- van Kerkwijk M. H., Kaplan D. L., 2007, *Ap&SS*, 308, 191
- van Straten W., Bailes M., 2011, *Proc. Astron. Soc. Aust.*, 28, 1
- Véron-Cetty M. P., Véron P., 2006, *A&A*, 455, 773
- Viganò D., Rea N., Pons J. A., Perna R., Aguilera D. N., Miralles J. A., 2013, *MNRAS*, 434, 123
- Wang H. G. et al., 2014, *ApJ*, 789, 73
- Wardle J. F. C., Kronberg P. P., 1974, *ApJ*, 194, 249
- Weltevrede P., 2016, *A&A*, 590, A109
- Weltevrede P., Johnston S., 2008, *MNRAS*, 391, 1210
- Young M. D., Manchester R. N., Johnston S., 1999, *Nature*, 400, 848
- Zarka P. et al., 2020, URSI GASS 2020, Session J01 NewTelescopes on the Frontier. Available at: <https://tinyurl.com/ycocd5ly>
- Zhang B., Harding A. K., Muslimov A. G., 2000, *ApJ*, 531, L135





**Figure A1.** The goodness-of-fit (reduced- $\chi^2$ ) of the RVM to the PA curve as a function of  $(\alpha, \beta)$  space obtained for the LOFAR Core data is shown in grey-scale. The black contours correspond to a reduced- $\chi^2$  of two-, three-, and four-times the minimum value. The green shaded regions are the ‘allowed’ viewing geometries, which are constraints arising from the estimated emission height and observed profile width (see the text in Appendix A for details including the assumptions made).

## APPENDIX A: CONSTRAINTS ON THE VIEWING GEOMETRY

This appendix gives a more detailed derivation of the viewing geometry based on fitting the RVM and considering the observed pulse width, which was summarized in Section 3.2. The RVM describes the shape of the PA ( $\psi$ ) curve as a function of pulse longitude,  $\phi$ . It depends on the magnetic inclination angle,  $\alpha$  and the impact parameter of the observer’s line of sight,  $\beta$ , and can be expressed as

$$\Delta\psi = \arctan\left(\frac{\sin(\Delta\phi)\sin\alpha}{\sin\zeta\cos\alpha - \cos\zeta\sin\alpha\cos(\Delta\phi)}\right), \quad (\text{A1})$$

where  $\Delta\psi = \psi - \psi_0$ ,  $\Delta\phi = \phi - \phi_0$ , and  $\zeta = \alpha + \beta$ . It describes a monotonic S-shaped curve which has an inflection point at  $(\phi_0, \psi_0)$ .

The goodness-of-fit of equation (A1) to the observed  $\psi(\phi)$  is parametrized by the reduced- $\chi^2$  and its variation is shown in Fig. A1 for the LOFAR Core data (see Rookyard et al. 2015b for details of the methodology used). The darker shading corresponds to lower reduced- $\chi^2$  values and so a better fit. The black contours indicate  $1\sigma$ ,  $2\sigma$  and  $3\sigma$  confidence intervals. As can be expected for a pulsar with a very small duty cycle,  $\alpha$  and  $\beta$  are highly correlated. The fit confirms that  $\beta$  must be small ( $<1.8^\circ$ ); however, the magnetic inclination  $\alpha$  is unconstrained from RVM fitting alone.

The measured profile widths provide additional information about the opening angle of the radio beam, how the line of sight cuts it, and the emission height. We assume that all radiation of a given frequency is produced at some height  $h_{\text{em}}$  in the magnetosphere in a circular region surrounding the magnetic axis. The emission beam is delimited by tangents to the last open field lines, forming a conal beam. In the small angle limit ( $h_{\text{em}} \ll R_{\text{LC}}$ ; see e.g. Rankin 1990),

the half opening angle of the emission cone is

$$\rho = \sqrt{\frac{9\pi h_{\text{em}}}{2cP}}. \quad (\text{A2})$$

This implies that the radio beam should widen with increasing emission height, and longer period pulsars can be expected to have narrower beams.

The width of the pulse profile depends on how the line of sight cuts through the emission beam. Gil, Gronkowski & Rudnicki (1984) showed that the rotational phase range for which the line of sight samples the open-field-line-region,  $W$ , can be expressed as

$$\cos\rho = \cos\alpha\cos(\alpha + \beta) + \sin\alpha\sin(\alpha + \beta)\cos\left(\frac{W}{2}\right). \quad (\text{A3})$$

This means that a measurement of  $W$  can help to constrain the parameters  $\alpha$  and  $\beta$ , as well as  $h_{\text{em}}$  via  $\rho$  (see e.g. Rookyard et al. 2015b). Here, it is important to note that the open-field-line region does not necessarily emit over its full extent; hence, the measured profile width does not necessarily correspond to  $W$  as defined in equation (A3).

The FAST profile is likely to correspond to a more fully illuminated beam (see Section 4.2). Therefore, Fig. A1 highlights the geometries which are compatible with the observed pulse width for that observation shown as the green shaded region. Here  $W_{10} = 4.3 \pm 0.2^\circ$  is used, the width of the profile as defined at 10 per cent of the peak flux density, to ensure that most emission is from the open-field-line region. Here it is assumed that the emission height lies within the range of 200–400 km (e.g. Mitra & Rankin 2002; Johnston & Karastergiou 2019). Moreover,  $W$  as defined in equation (A3) is assumed to be between the measured  $W_{10}$  and twice the distance between the PA curve inflection point and the furthest edge of the FAST pulse profile, in order to account for potential underfilling of the radio beam (see Section 4.3 for the motivation). Here, we take the PA inflection point to coincide with the position of the fiducial plane (the plane containing the magnetic and rotation axes), because the emission height at FAST frequencies is argued to be low enough to make any A/R effects small (see Section 4.3). Taking into account the uncertainties on  $W_{10}$  and the inflection point, this results in  $4.1^\circ \leq W \leq 6.9^\circ$  assuming that at least one edge of the profile corresponds to the boundary with the last open-field-line region. This allowed range of  $W$  and emission height results in a collection of contours in  $(\alpha, \beta)$  space defined by equations (A2) and (A3). These contours (green shaded regions Fig. A1) show that  $\beta$  is likely  $\leq 1.1^\circ$ , and suggests that the pulsar is relatively aligned (a small  $\alpha$ ). Further to these considerations, it cannot be ruled out that neither edge of the profile reaches the edge of the open-field-line region. This would correspond to a lower filling fraction and modestly more aligned  $\alpha$ .

## APPENDIX B: THE CORE-CONE MODEL

The profile width evolution of PSR J0250+5854 can be described by the core-cone model, and the geometry derived in this way is consistent with our results in Section 3.2 and Appendix A. In this scenario, the LOFAR Core profile is associated with emission from the core beam, while the FAST profile is broader due to presence of emission from the inner cone, in the form of ‘conal outriders’. The width of the core beam is believed to reflect the size of the polar cap at the surface of the pulsar, and is largely independent of observing frequency (Rankin 1983a) and follows the empirical relation  $W_{50} = 2.45^\circ P^{-0.5}/\sin\alpha$  (Rankin 1990). For PSR J0250+5854,  $P = 23.5$  s

and the measured width of the LOFAR Core profile  $W_{50} \approx 1^\circ$  gives  $\alpha \approx 30^\circ$ . The gradient of the PA curve at the inflection point  $\sin \alpha / \sin \beta \approx 55 \text{ deg deg}^{-1}$  which then implies  $\beta \approx 0.5^\circ$ . These values of  $\alpha$  and  $\beta$  lie within the ‘allowed geometries’ in Fig. A1, which were derived without a specific model in mind.

Extending this further, Rankin (1993) argued that the half opening angle of the inner cone follows the relation  $\rho_{\text{inner}} = 4.3^\circ P^{-0.52}$ , which for PSR J0250+5854 gives  $\rho_{\text{inner}} = 0.83^\circ$ . Substituting this and the derived  $\alpha$  and  $\beta$  values into equation (A3) gives the expected width of the profile due to the inner cone,  $W \approx 2.5^\circ$ . This is consistent with  $W_{50} = 2.4 \pm 0.1^\circ$  measured for the FAST profile.

<sup>1</sup>Jodrell Bank Centre for Astrophysics, School of Physics and Astronomy, University of Manchester, Manchester M13 9PL, UK

<sup>2</sup>Laboratoire de Physique et Chimie de l’Environnement et de l’Espace (LPC2E), Université d’Orléans/CNRS, F-45071 Orléans, France

<sup>3</sup>LESIA, Observatoire de Paris, CNRS, PSL, SU/UP/VO, F-92195 Meudon, France

<sup>4</sup>Station de Radioastronomie de Nançay, Observatoire de Paris, PSL Research University, CNRS, Univ. Orléans, OSUC, F-18330 Nançay, France

<sup>5</sup>ASTRON, Netherlands Institute for Radio Astronomy, Oude Hoogeveensedijk 4, NL-7991 PD Dwingeloo, the Netherlands

<sup>6</sup>Anton Pannekoek Institute for Astronomy, University of Amsterdam, Science Park 904, NL-1098 XH, Amsterdam, the Netherlands

<sup>7</sup>Department of Astronomy, School of Physics and Materials Science, Guangzhou University, Guangzhou 510006, China

<sup>8</sup>Astronomical Science and Technology Research Laboratory, Department of Education of Guangdong Province, Guangzhou 510006, China

<sup>9</sup>Xinjiang Astronomical Observatory, Chinese Academy of Sciences, 150 Science 1-Street, Urumqi, Xinjiang 830011, China

<sup>10</sup>Astrophysics, Department of Physics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford, OX1 3RH, UK

<sup>11</sup>Astro Space Centre, Lebedev Physical Institute, Russian Academy of Sciences, Profsoyuznaya Str 84/32, Moscow 117997, Russia

<sup>12</sup>Fakultät für Physik, Universität Bielefeld, Postfach 100131, D-33501 Bielefeld, Germany

<sup>13</sup>CAS Key Laboratory of FAST, National Astronomical Observatories, Chinese Academy of Sciences, Beijing 100101, China

<sup>14</sup>University of Chinese Academy of Sciences, Beijing 100049, China

<sup>15</sup>NAOC-UKZN Computational Astrophysics Centre, University of KwaZulu-Natal, Durban 4000, South Africa

<sup>16</sup>CSIRO Astronomy and Space Science, PO Box 1130, Bentley, WA 6102, Australia

<sup>17</sup>Department of Physics, McGill University, 3600 Rue University, Montréal, QC H3A 2T8, Canada

<sup>18</sup>McGill Space Institute, McGill University, 3550 Rue University, Montréal, QC H3A 2A7, Canada

<sup>19</sup>Laboratoire Univers et Théories, Observatoire de Paris, Université PSL, CNRS, Université de Paris, F-92190 Meudon, France

<sup>20</sup>AIM, CEA, CNRS, Université Paris-Saclay, Université de Paris, F-91191 Gif-sur-Yvette, France

<sup>21</sup>Sorbonne Université, Université de Paris, 5 Place Jules Janssen, F-92195 Meudon, France

This paper has been typeset from a  $\text{\TeX}/\text{\LaTeX}$  file prepared by the author.