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Pavel Mancera Piña, Lorenzo Posti, Gabriele Pezzulli, Filippo Fraternali, S. Michael Fall, et al.. A tight angular-momentum plane for disc galaxies. *Astronomy and Astrophysics - A&A*, 2021, 651, pp.L15. 10.1051/0004-6361/202141574 . hal-03337177

HAL Id: hal-03337177

<https://hal.science/hal-03337177>

Submitted on 2 Aug 2022

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LETTER TO THE EDITOR

A tight angular-momentum plane for disc galaxies[★]

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Received 16 June 2021 / Accepted 5 July 2021

ABSTRACT

The relations between the specific angular momenta (j) and masses (M) of galaxies are often used as a benchmark in analytic models and hydrodynamical simulations as they are considered to be amongst the most fundamental scaling relations. Using accurate measurements of the stellar (j_*), gas (j_{gas}), and baryonic (j_{bar}) specific angular momenta for a large sample of disc galaxies, we report the discovery of tight correlations between j , M , and the cold gas fraction of the interstellar medium (f_{gas}). At fixed f_{gas} , galaxies follow parallel power laws in 2D (j, M) spaces, with gas-rich galaxies having a larger j_* and j_{bar} (but a lower j_{gas}) than gas-poor ones. The slopes of the relations have a value around 0.7. These new relations are amongst the tightest known scaling laws for galaxies. In particular, the baryonic relation ($j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$), arguably the most fundamental of the three, is followed not only by typical discs but also by galaxies with extreme properties, such as size and gas content, and by galaxies previously claimed to be outliers of the standard 2D $j - M$ relations. The stellar relation ($j_* - M_* - f_{\text{gas}}$) may be connected to the known $j_* - M_*$ -bulge fraction relation; however, we argue that the $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ relation can originate from the radial variation in the star formation efficiency in galaxies, although it is not explained by current disc instability models.

Key words. galaxies: kinematics and dynamics – galaxies: spiral – galaxies: dwarf – galaxies: formation – galaxies: evolution – galaxies: fundamental parameters

1. Introduction

Despite the remarkable diversity of galaxy properties observed in the present-day Universe, a number of physical parameters of galaxies appear to correlate with one another and form tight scaling laws. Such relations are of paramount importance in our quest to understand galaxy formation and evolution (e.g., Tully & Fisher 1977; Fall 1983; Burstein et al. 1997; McGaugh et al. 2000; Marconi & Hunt 2003; Cappellari et al. 2013; Wang et al. 2016).

Since early models of galaxy formation were proposed, it has become clear that mass and angular momentum are two fundamental parameters controlling the evolution of galaxies (e.g., Fall & Efstathiou 1980; Dalcanton et al. 1997; Mo et al. 1998). From an observational point of view, starting from the work by Fall (1983), different authors have characterised the scaling relation between stellar mass (M_*) and stellar specific angular momentum ($j_* = J_*/M_*$, where J_* is the angular momentum) as the $j_* - M_*$ or Fall relation (e.g., Romanowsky & Fall 2012; Obreschcow & Glazebrook 2014; Posti et al. 2018a). This $j_* - M_*$ law has been widely used in recent years to constrain and test both (semi-)analytic models and hydrodynamical simulations (e.g., Genel et al. 2015; Pedrosa & Tissera 2015; Obreja et al. 2016; Lagos et al. 2017; Tremmel et al. 2017; El-Badry et al.

2018; Stevens et al. 2018; Zoldan et al. 2018; Irodotou et al. 2019).

Mancera Piña et al. (2021, hereafter MP21), derived accurate measurements of the stellar (j_*), (cold neutral) gas (j_{gas}), and baryonic (j_{bar}) specific angular momentum for a large sample of irregular spiral and dwarf galaxies (see also e.g., Obreschcow & Glazebrook 2014; Kurapati et al. 2018). They determined the $j - M$ relations for the three components and fitted them with unbroken power laws. They also noticed that the residuals from the best fitting relations correlate with the gas fraction ($f_{\text{gas}} = M_{\text{gas}}/M_{\text{bar}}$, with M_{gas} and M_{bar} the gas and baryonic masses, respectively). These trends, also seen in a few semi-analytic models (e.g., Stevens et al. 2018; Zoldan et al. 2018), may indicate that the gas content plays an important role in the $j - M$ relations. In this Letter we build upon that result and report the discovery of new and very tight correlations between mass, specific angular momentum, and gas fraction. We show that disc galaxies across ~ 4 orders of magnitude in mass lie in very tight planes in the (j, M, f_{gas}) spaces.

2. Definition of j and galaxy sample

The stellar and gas specific angular momenta of a galaxy are defined as

$$j_i(< R) = \frac{\int_0^R R'^2 \Sigma_i(R') V_i(R') dR'}{\int_0^R R' \Sigma_i(R') dR'} \quad (1)$$

[★] The associated data catalogue is only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/cat/J/A+A/651/L15>

with R being the galactocentric cylindrical radius, Σ_i the stellar or gas face-on surface density, and V_i the stellar or gas rotation velocity. Then, j_* and j_{gas} can be combined to obtain

$$j_{\text{bar}} = f_{\text{gas}} j_{\text{gas}} + (1 - f_{\text{gas}}) j_*. \quad (2)$$

For $f_{\text{gas}} = M_{\text{gas}}/M_{\text{bar}}$, we assumed $M_{\text{bar}} = M_* + M_{\text{gas}}$, with $M_{\text{gas}} = 1.33M_{\text{HI}}$, where M_{HI} is the mass of neutral atomic hydrogen and the factor 1.33 accounts for the presence of helium. While we neglected any contribution from molecular gas, in Appendix A we show that its inclusion does not change the results found in this Letter¹.

MP21 compiled a high-quality sample of 157 nearby galaxies, predominantly discs. All the galaxies have near-IR photometry and extended HI rotation curves, allowing their stellar discs and rotation velocities to be traced robustly. The sample includes dwarf and massive galaxies, spanning the mass range $7 \lesssim \log(M_*/M_\odot) \lesssim 11.5$, $6 \lesssim \log(M_{\text{gas}}/M_\odot) \lesssim 10.5$, $8 \lesssim \log(M_{\text{bar}}/M_\odot) \lesssim 11.5$, and with $0.01 < f_{\text{gas}} < 0.97$ and a typical relative uncertainty, $\delta f_{\text{gas}}/f_{\text{gas}} \approx 0.2$ dex (median $\delta f_{\text{gas}} \approx 0.05$). While not complete, the sample is representative of the population of regularly rotating nearby discs, like other large samples commonly used in the literature (e.g., Lelli et al. 2016; Ponomareva et al. 2016). Using the near-IR photometry and HI rotation curves, MP21 built cumulative radial profiles for j_* (applying a correction to convert V_{gas} into V_*), j_{gas} , and j_{bar} . By selecting only galaxies with radially convergent measurements of angular momentum, they built a sample of 130, 87, and 106 galaxies with accurate j_* , j_{gas} , and j_{bar} , respectively. For more details we refer the reader to MP21, and in our associated data catalogues we provide the values of j , M , f_{gas} , distance, and Hubble type for the galaxy sample. These tables are available at the CDS.

MP21 fitted the $j - M$ relations with power laws of the form

$$\log\left(\frac{j_i}{\text{kpc km s}^{-1}}\right) = m_i[\log(M_i/M_\odot) - 10] + n_i, \quad (3)$$

with the subscript i representing the stellar, gaseous, or baryonic component. The best fitting power laws have slopes m_i of about 0.5, 1.0, and 0.6 for stars, gas, and baryons, respectively, and an intrinsic scatter of 0.15 dex.

3. The $j - M - f_{\text{gas}}$ planes

3.1. Best fitting planes

MP21 (see their Fig. 7) also found systematic trends with f_{gas} in the residuals of the three $j - M$ laws. To see if introducing a dependence of the $j - M$ laws on f_{gas} can explain these trends, we fitted the (j, M, f_{gas}) data with planes. We fitted the data points with the model

$$\log\left(\frac{j_i}{\text{kpc km s}^{-1}}\right) = \alpha_i \log(M_i/M_\odot) + \beta_i \log(f_{\text{gas}}) + \gamma_i. \quad (4)$$

¹ Our ‘gas’ refers only to the interstellar medium, and there is no attempt to include the largely unconstrained contribution of the gas outside galaxy discs. Although the sum of our gas and stars does not represent the ‘whole’ baryonic budget of a galaxy, we prefer to keep this nomenclature for the sake of consistency with the literature.

Table 1. Coefficients of the best fitting $j - M - f_{\text{gas}}$ planes.

	α	β	γ	σ_\perp
Stars	0.67 ± 0.03	0.51 ± 0.08	-3.62 ± 0.23	0.10 ± 0.01
Gas	0.78 ± 0.03	-0.49 ± 0.04	-4.64 ± 0.25	0.08 ± 0.01
Baryons	0.73 ± 0.02	0.46 ± 0.05	-4.25 ± 0.19	0.08 ± 0.01

Therefore, we assumed that, in contrast to Eq. (3), j_i also depends on f_{gas} ². We performed the fit using the R package HYPER-FIT (Robotham & Obreschkow 2015), including a term for the intrinsic scatter σ_\perp . We assumed log-normal uncertainties in j , M , and f_{gas} , and, using a Monte Carlo sampling method, we took into account the fact that uncertainties in the distance, inclination, and mass-to-light ratio of a given galaxy drive correlated uncertainties (also provided in our electronic tables) between $\log(M)$, $\log(j)$, and $\log(f_{\text{gas}})$. We stress that taking these correlations into account is important: Neglecting them can artificially lower the intrinsic scatter of the planes by a factor of two to three.

The best fitting coefficients are reported in Table 1. The orthogonal intrinsic scatter of our best fitting planes is significantly smaller than for the 2D relations. The log-marginal likelihood is also higher (i.e., better) for the 3D planes: by 27, 40, and 43 units for stars, gas, and baryons, respectively. We conclude that the inclusion of f_{gas} into the $j - M$ laws is statistically meaningful.

In Fig. 1 we compare the observed distribution of galaxies with our three best fitting planes; the figure shows the 3D $j - M - f_{\text{gas}}$ planes projected into the 2D (j, M) spaces. Galaxies are colour-coded according to their f_{gas} , and we overlay our lines of constant f_{gas} derived from our best fitting planes. The fits provide a very good description of the data, in line with the low intrinsic scatter we find for all the planes. By construction, the three $j - M - f_{\text{gas}}$ planes are characterised by their M slopes (α_* , α_{gas} , and α_{bar}) and f_{gas} slopes (β_* , β_{gas} , and β_{bar}). Projected into the (j, M) spaces, the f_{gas} slopes act as a normalisation for j . At fixed M_* (M_{bar}), gas-rich galaxies have a higher j_* (j_{bar}) than gas-poor ones, while gas-poor galaxies show a higher j_{gas} . For stars and baryons, the 3D relations become steeper ($\alpha_* = 0.67 \pm 0.03$, $\alpha_{\text{bar}} = 0.73 \pm 0.03$) than the 2D ones from MP21 ($m_* = 0.53 \pm 0.02$, $m_{\text{bar}} = 0.60 \pm 0.02$) once f_{gas} is taken into account, while the slope of the gas relation becomes shallower ($\alpha_{\text{gas}} = 0.78 \pm 0.03$, $m_{\text{gas}} = 1.02 \pm 0.04$). Given the different coefficients, the 2D $j - M$ relations (shown in Fig. 1 as green bands) differ from the projection of the 3D planes in the (j, M) spaces, especially at $M < 10^8 M_\odot$.

A remarkable property of our new scaling laws is their low intrinsic scatter. Given that the baryonic $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ plane incorporates the stellar and cold gas components, we argue that this is likely the most fundamental of the three relations, although its intrinsic scatter is similar to the relation for the gas. Very few other scaling laws are thought to have a comparably low intrinsic scatter, for instance the HI mass-size relation (Wang et al. 2016) or the baryonic Tully-Fisher relation (BTFR; McGaugh et al. 2000; Ponomareva et al. 2017). In fact, our bary-

² Since our planes depend on $\log(f_{\text{gas}})$, they become hard to interpret when $f_{\text{gas}} \rightarrow 0$, preventing us from making extrapolations for galaxies with $f_{\text{gas}} < 0.01$. Using f_{gas} instead of $\log(f_{\text{gas}})$ produces less satisfactory fits when compared to the observations, so we prefer $\log(f_{\text{gas}})$ despite its limitations when $f_{\text{gas}} \rightarrow 0$. We also note that, given Eq. (2), the $j - M$ relations cannot all be exactly planes. However, fitting planes in the (j, M, f_{gas}) spaces is a very useful empirical approach.

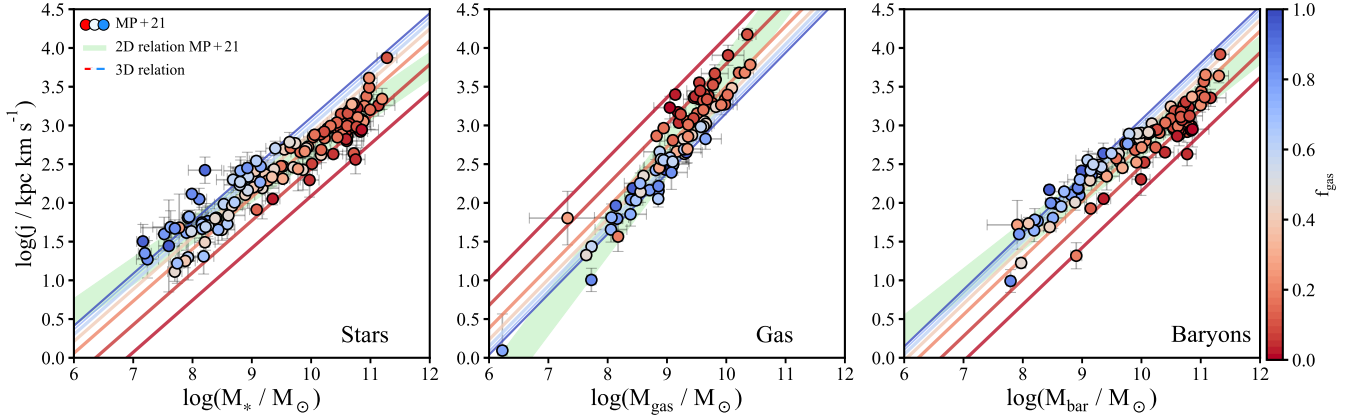


Fig. 1. Stellar, gas, and baryonic $j - M - f_{\text{gas}}$ planes, projected into the 2D (j, M) spaces. Galaxies are colour-coded according to their f_{gas} and are compared with lines of constant f_{gas} according to Eq. (4) and the best fitting coefficients of Table 1. From red to blue, the lines are at $f_{\text{gas}} = 0.01, 0.05, 0.2, 0.4, 0.6, 0.8, 1$. For comparison, we show in green the best fitting 2D $j - M$ relations from MP21 and their intrinsic scatter.

onic plane can in principle be used as a distance estimator, with an uncertainty $\delta D/D = (\delta j_{\text{bar}}/j_{\text{bar}})/|2\alpha_{\text{bar}} - 1|$ at fixed M_{bar} .

3.2. The similarities of the α slopes

The three α slopes of our $j - M - f_{\text{gas}}$ planes are relatively close to one another and to the value $2/3$ expected for their parent dark matter halos (Fall 1983), which suggests some degree of structural self-similarity between different baryonic components and the dark matter halo. From a mathematical point of view, if we rewrite Eq. (2) in terms of $j_* = BM_*^{\alpha_*}$ (with B a function that depends only on f_{gas}) and $M_* = M_{\text{bar}}(1 - f_{\text{gas}})$, we obtain

$$j_{\text{bar}} = B(1 - f_{\text{gas}})^{\alpha_*} \left[\frac{j_{\text{gas}}}{j_*} f_{\text{gas}} + (1 - f_{\text{gas}}) \right] M_{\text{bar}}^{\alpha_*}. \quad (5)$$

In a similar way, considering now $j_{\text{gas}} = CM_{\text{gas}}^{\alpha_{\text{gas}}}$ (with C a function that depends only on f_{gas}), we find

$$j_{\text{bar}} = C f_{\text{gas}}^{\alpha_{\text{gas}}} \left[\frac{j_*}{j_{\text{gas}}} (1 - f_{\text{gas}}) + f_{\text{gas}} \right] M_{\text{bar}}^{\alpha_{\text{gas}}}. \quad (6)$$

Therefore, at fixed f_{gas} , the slope α_{bar} of the baryonic $j - M - f_{\text{gas}}$ plane is expected to be similar to α_* and α_{gas} , provided that the ratio j_{gas}/j_* is independent of M_{bar} . As shown in Fig. 2, for our sample, j_{gas}/j_* , which is always larger than 1 and mostly within a narrow range (the 16th and 84th percentiles are 1.5 and 3.2, respectively), does not seem to correlate with M_{bar} , in line with the near-parallelism of the three relations shown in Fig. 1.

It is worth noticing that the j_{gas}/j_* ratio can be related to the relative extent of some characteristic size of the gaseous (R_{gas}) and stellar (R_*) components of galactic discs, given that $j_{\text{gas}}/j_* \approx R_{\text{gas}} V_{\text{gas}} / (R_* V_*) \approx R_{\text{gas}} / R_*$. Although this is just an approximation, it can be useful in the physical interpretation of the $j - M - f_{\text{gas}}$ relations (e.g., Sect. 4.2.2). In Appendix B we show the expected dependence of j_{gas}/j_* on M_{bar} and f_{gas} derived from our best fitting $j_* - M_* - f_{\text{gas}}$ and $j_{\text{gas}} - M_{\text{gas}} - f_{\text{gas}}$ planes.

3.3. No outliers of the baryonic $j - M - f_{\text{gas}}$ law

In Fig. 3 we plot again our baryonic plane, this time splitting the galaxies into bins of f_{gas} . In each panel we plot lines covering the whole range of f_{gas} within that bin. This allows the tightness of our baryonic relation to be appreciated in more detail.

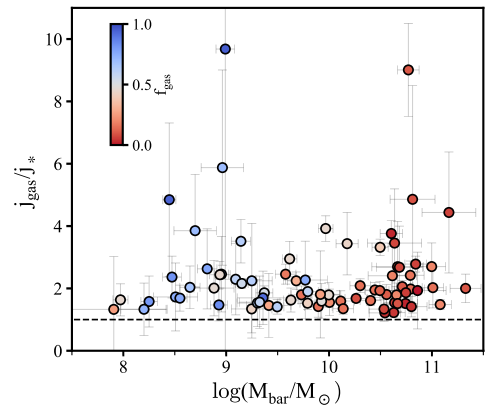


Fig. 2. j_{gas}/j_* ratio as a function of M_{bar} . Galaxies are colour-coded by their f_{gas} , and the dashed black line corresponds to $j_{\text{gas}}/j_* = 1$. Our galaxies (those with convergent j_* and j_{gas} from MP21) cluster at $j_{\text{gas}}/j_* \sim 2$ at all M_{bar} , albeit with a significant scatter.

We now investigate whether any objects from known galaxy populations could be outliers of our baryonic plane. We do this by comparing our relation – derived using only our sample – with galaxies from the literature that have been argued to be outliers of the 2D $j - M$ relations or that could be outliers given their extreme properties in size, f_{gas} , or rotation velocity.

First, we considered a set of dwarf galaxies from the literature, specifically those from Elson (2017) and Kurapati et al. (2018), that do not overlap with our sample. It was claimed by those authors that some of these galaxies are off the 2D $j_{\text{bar}} - M_{\text{bar}}$ relation. We also included the dwarf ‘super thin’ galaxies of Jadhav & Banerjee (2019), which have very high axis ratios and have been suggested to have higher j_* than other dwarfs. Next, we looked at the ‘HI extreme galaxies’ (HIXs) of Lutz et al. (2018), which have a particularly high f_{gas} for their M_* and are claimed to have higher j_{bar} than average. Moreover, we added a sample of ‘super spiral’ galaxies (di Teodoro et al. 2021), which are very large discs with masses a factor of ten larger than L_* galaxies and were also claimed to be outliers of the $j_* - M_*$ relation (Ogle et al. 2019)³. Finally, we included the

³ Since j_{gas} is not available for the super spirals, we assumed $j_{\text{gas}} = 1.9j_*$, the median ratio found in the rest of our sample. This has little to no impact given their low gas content ($f_{\text{gas}} \approx 0.15$).

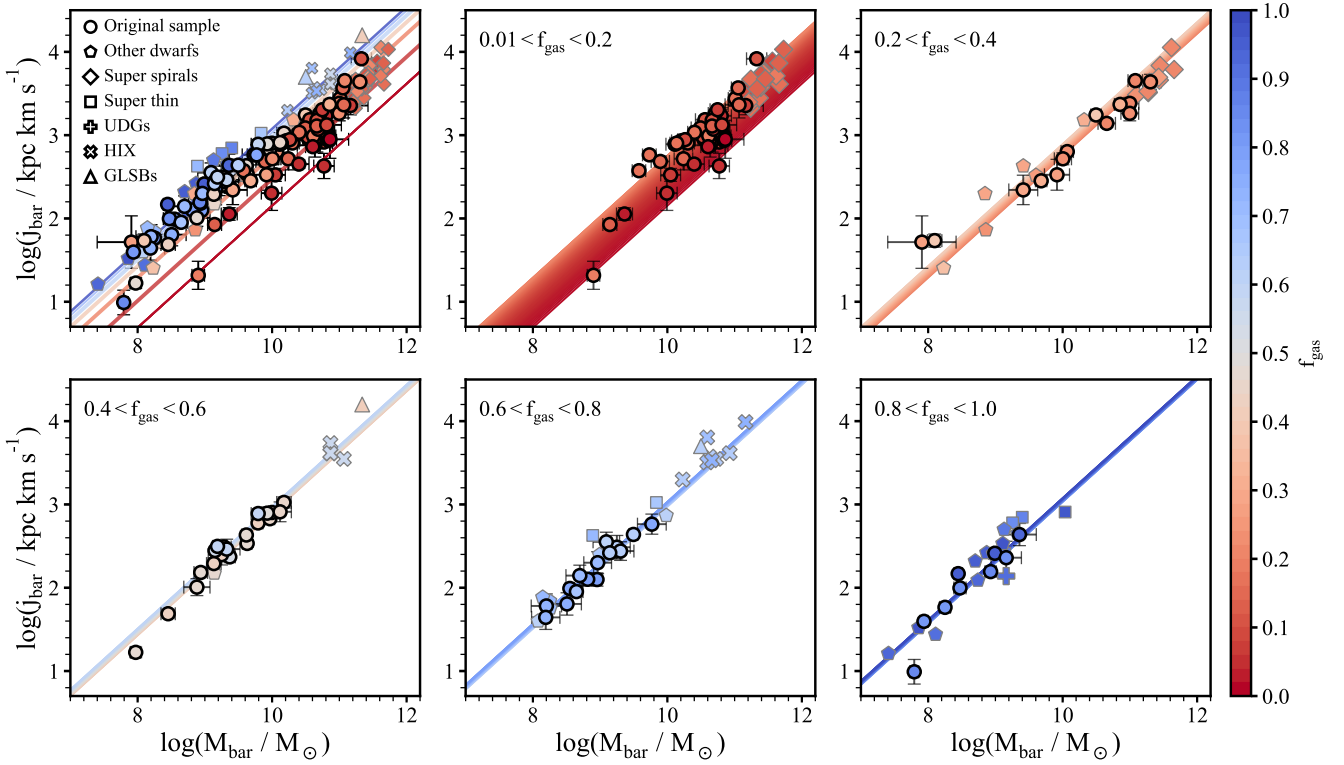


Fig. 3. Baryonic $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ plane for our original sample (circles) and a set of extreme galaxies (see text). *Top left panel:* relation for all the galaxies, while the remaining panels show the galaxies in bins of f_{gas} (given in the *top left corner* of each panel). *First panel,* the lines of constant f_{gas} are as in Fig. 1. In the remaining panels, the coloured areas enclose the region delimited by the whole f_{gas} bin. We remark that the coloured lines of our plane are derived by fitting only our data. The rest of the galaxies closely follow our fit.

ultra-diffuse galaxies (UDGs) AGC 114905 and AGC 242019 (Mancera Piña et al. in prep.; Shi et al. 2021) and the giant low surface brightness galaxies (GLSBs) Malin 1 and NGC 7589 (Lelli et al. 2010). The UDGs are found to be outliers of the BTFR (Mancera Piña et al. 2019, 2020), and both UDGs and GLSBs are extreme galaxies with very extended discs for their M_* . With the caveat that some of these galaxies have larger uncertainties than our sample, given the different data quality, we added all these sets of galaxies into the $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ plane in Fig. 3. Remarkably, the location of all of these galaxies, given their f_{gas} , is in very good agreement with the expectation of our scaling relation. We conclude that even extreme galaxies such as HIXs, UDGs, and GLSBs obey the $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ law.

4. Discussion

4.1. Stellar relation: The link with bulge fraction

Previous works (e.g., Fall 1983; Romanowsky & Fall 2012; Cortese et al. 2016; Fall & Romanowsky 2018) found that the relation between j_* and M_* depends on the bulge-to-total mass fraction, \mathcal{B}_* . Fall & Romanowsky (2018, hereafter FR18), proposed a model in which discs and spheroids follow relations of the form $j_* \propto M_*^{0.67}$ but with different normalisation, with spheroids shifting downwards with respect to discs; any given galaxy then has a j_* that can be expressed as a linear superposition of a disc and a spheroid (a bulge).

The resemblance of our $j_* - M_* - f_{\text{gas}}$ plane (where at fixed M_* a different f_{gas} produces a shift in j_*) with the \mathcal{B}_* relation is clear. Both relations are valid for a variety of morphological types and are preserved along a broad mass span, and with a dependence of j_* on M_* with a slope of 2/3. The similarities are not

unexpected since gas-poor galaxies usually have high \mathcal{B}_* , though the $f_{\text{gas}} - \mathcal{B}_*$ relation is highly scattered. The above suggests that these two relations may be manifestations of a common physical mechanism. Finally, we note that the scatter is better quantified in the $j_* - M_* - f_{\text{gas}}$ relation with respect to the \mathcal{B}_* relation given that the uncertainties in \mathcal{B}_* are difficult to estimate (Salo et al. 2015; FR18).

Interestingly, there is a regime in which the f_{gas} relation makes a different prediction from the \mathcal{B}_* relation. For galaxies that are almost gas-free and almost bulge-less (we note that galaxies with $\mathcal{B}_* \lesssim 0.1 - 0.2$ host pseudo-bulges rather than classical bulges; see Fig. 3 in FR18), the f_{gas} relation expects them to have a lower-than-average j_* , while the \mathcal{B}_* relation predicts a higher-than-average j_* . We tested this in Fig. 4 by looking at the four galaxies in our sample with $f_{\text{gas}} \leq 0.1$ and $\mathcal{B}_* \leq 0.1$ (Salo et al. 2015). We also plot the expected lines given the average \mathcal{B}_* and f_{gas} for these four galaxies. The points lie close to the f_{gas} relation and deviate from the \mathcal{B}_* one. However, the evidence is not compelling given the low-number statistics. Finally, it is also important to mention that our galaxy sample is fairly different from that of FR18, with many more gas-dominated discs but a lack of early-type galaxies. These potential differences can be further explored with larger and more complete samples.

4.2. The origin of the baryonic relation

We then focused our attention on the $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ relation. Its origin is likely related to different galaxy formation processes, such as variations in the angular momentum of the dark matter halos, selective gas accretion within the discs, and different gas accretion histories from the intergalactic medium (Fall 1983; Posti et al. 2018a,b; Stevens et al. 2018; Zoldan et al. 2018).

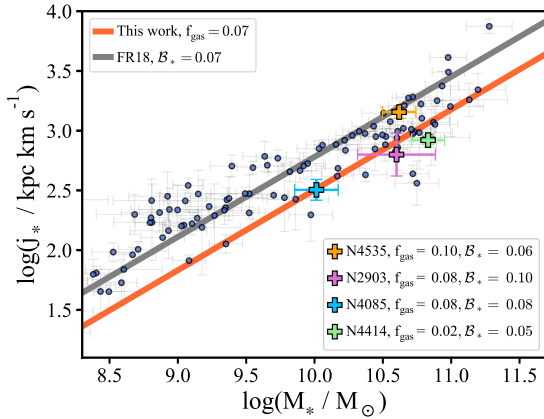


Fig. 4. $j_* - M_*$ relation. Crosses show our galaxies with $(f_{\text{gas}}, \mathcal{B}_*) \leq 0.1$. The black and red lines show the expectations from FR18 and this work, respectively. Only a fraction of our range in M_* is shown.

Evolutionary processes such as stellar and active galactic nucleus feedback, mergers, and angular momentum transfer between galaxies and their dark matter halos are arguably also important (Leroy et al. 2008; Dutton & van den Bosch 2012; Romanowsky & Fall 2012; Lagos et al. 2017; Zoldan et al. 2018). Still, while a complex interplay between all these processes is expected, it all results in the tight $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ law we observe. Therefore, it is interesting to check whether or not simple mechanisms are able to capture the dominant processes that give rise to the $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ relation.

4.2.1. Disc instability

We consider two models that have been proposed to explain the $j_{\text{bar}} - f_{\text{gas}}$ connection as a consequence of gravitational instability. First, Obreschkow et al. (2016) proposed a model that relates f_{gas} with j_{bar} via

$$f_{\text{gas}} = \min\{1, 2.5q^{1.12}\}, \quad q = j_{\text{bar}}\sigma_{\text{gas}}/GM_{\text{bar}}, \quad (7)$$

with q a stability parameter, σ_{gas} the gas velocity dispersion, and G the gravitational constant. Deviations from Eq. (7) may occur depending on the galaxy rotation curve, but they are expected to be small. These results were derived under a number of simplifying assumptions, but less idealised semi-analytic models are found to be in good agreement (Stevens et al. 2018). From Eq. (7), one has $\log(f_{\text{gas}}) \propto 1.12[\log(j_{\text{bar}}) - \log(M_{\text{bar}})]$ and $j_{\text{bar}} \propto M_{\text{bar}}$ (this at fixed f_{gas} and assuming a constant σ_{gas}). These dependences disagree with our best fitting plane, which has $\log(f_{\text{gas}}) \propto 2.17 \log(j_{\text{bar}}) - 1.59 \log(M_{\text{bar}})$ and $j_{\text{bar}} \propto M_{\text{bar}}^{0.73}$ at fixed f_{gas} . Projecting our baryonic plane into the $f_{\text{gas}} - q$ diagram shows that galaxies of a given M_{bar} follow parallel sequences of the form $f_{\text{gas}} \propto q^{1/\beta_{\text{bar}}} = q^{2.22}$, instead of $f_{\text{gas}} \propto q^{1.12}$ ⁴.

Also based on disc instability, Romeo (2020) proposed a set of scaling relations of the form $j_i \hat{\sigma}_i / (GM_i) \approx 1$, with i denoting stars or gas and $\hat{\sigma}$ a mass-weighted radial average of the velocity dispersion σ . This relation produces the scaling $j_* \propto M_*^{0.5}$ (for $\hat{\sigma}_* \propto M_*^{0.5}$, as proposed by Romeo 2020) and $j_{\text{gas}} \propto M_{\text{gas}}$, very similar to the values found in MP21 for the 2D relations. To incorporate f_{gas} , we rewrote the above expression (using

⁴ We note that assuming a non-constant σ_{gas} is not enough to alleviate the mentioned discrepancies. To match our relations, $\sigma_{\text{gas}} \propto M_{\text{bar}}^{-0.25}$ is required; however, it is observed that $\sigma_{\text{gas}} \propto M_{\text{bar}}^{0.07}$ (e.g., Murugesan et al. 2020).

$M_i = M_{\text{gas}} = f_{\text{gas}} M_{\text{bar}}$) as

$$\frac{j_{\text{gas}}}{M_{\text{bar}}} \frac{\hat{\sigma}_{\text{gas}}}{G} = f_{\text{gas}}. \quad (8)$$

Assuming a constant $\hat{\sigma}_{\text{gas}}$, as in Romeo (2020), this relation predicts $j_{\text{gas}} \propto M_{\text{bar}}$ at fixed f_{gas} and $j_{\text{gas}} \propto f_{\text{gas}}$ at fixed M_{bar} . Instead, a corollary of our gas relation is that $j_{\text{gas}} \propto M_{\text{bar}}^{0.78}$ at fixed f_{gas} and that $j_{\text{gas}} \propto f_{\text{gas}}^{0.27}$ at fixed M_{bar} . Therefore, the relation from Romeo (2020) also seems to depart from our data.

4.2.2. Star formation efficiency

A more general possibility is that the link between j_{bar} , M_{bar} , and f_{gas} is related to the star formation efficiency in galaxies, as we briefly discuss here. We started by noting that at fixed M_{bar} the larger the j_{bar} of a galaxy is, the more extended its baryonic distribution will be⁵. Also, it is well established that gas located in the outskirts of discs forms stars less efficiently than gas closer to the centre (e.g., Leroy et al. 2008). Thus, at fixed M_{bar} , a galaxy with a large j_{bar} also has a large f_{gas} since a large portion of its mass is located in the less star-forming outer regions. Qualitatively, this is in agreement with the fact that for our entire galaxy sample $j_{\text{gas}}/j_* \approx R_{\text{gas}}/R_* > 1$ (see Sect. 3.2). All this suggests that the connection between j_{bar} and f_{gas} may be a reflection of the mechanism responsible for a radially declining star formation efficiency and a radially increasing f_{gas} in galaxy discs (e.g., Leroy et al. 2008; Krumholz et al. 2011; Bacchini et al. 2019), but exploring this idea quantitatively (e.g., by investigating why the j_{gas}/j_* ratio is largely independent of M_{bar} and f_{gas} ; see Fig. 2 but also Fig. B.1) is beyond the scope of the present Letter.

5. Conclusions

In this Letter we have used a high-quality sample of disc galaxies to study the relation between their specific angular momenta (j), masses (M), and gas fraction (f_{gas}). The position of our galaxies in the (j, M, f_{gas}) spaces can be described with planes such that galaxies with different f_{gas} follow parallel lines in the projected 2D (j, M) spaces. Remarkably, our planes are preserved along a wide range of mass and morphology with very small (≤ 0.1 dex) intrinsic scatter, which places the relations amongst the tightest scaling laws for disc galaxies. The $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ plane is arguably the most fundamental, and it is even followed by populations of galaxies with extreme size, mass, and gas content, some of which were previously claimed to be outliers of the 2D $j - M$ relations.

The $j_* - M_* - f_{\text{gas}}$ relation shows analogies with the $j_* - M_* - \mathcal{B}_*$ relation (\mathcal{B}_* being the bulge-to-total mass fraction) previously discussed in the literature. Galaxies with fixed f_{gas} or \mathcal{B}_* follow parallel relations of the form $j_* \propto M_*^{2/3}$. Most galaxies are well described by both the f_{gas} and \mathcal{B}_* relations, while some discrepancies appear when looking at galaxies with low f_{gas} and low \mathcal{B}_* .

Finally, we show that models based purely on disc instability do not quantitatively reproduce the observed $j - M - f_{\text{gas}}$ relations. We argue that the origin and behaviour of the $j_{\text{bar}} - M_{\text{bar}} - f_{\text{gas}}$ law is closely related to the spatial distribution of gas and stars within galaxies, as well as to the radial variations in the star formation efficiency.

⁵ Second-order effects related to the concentration of the host halo might also be relevant in determining the galaxy sizes (Posti et al. 2020).

We stress that our relations offer a unique possibility to quantitatively test a variety of models, including hydrodynamical simulations and semi-analytic models, providing a powerful benchmark for theories on the formation of galactic discs. The slopes and intrinsic scatter of our $j - M - f_{\text{gas}}$ planes are important requirements that hydrodynamical simulations and (semi-)analytic models should aim to reproduce.

Acknowledgements. We would like to thank our referee, Danail Obreschkow, for a very constructive and valuable report, and for his kind assistance regarding the use of HYPER-FIT. We thank Enrico di Teodoro for making available the data on ‘super spiral’ galaxies for us. P.E.M.P., and F.F. are supported by the Netherlands Research School for Astronomy (Nederlandse Onderzoekschool voor Astronomie, NOVA), Project 10.1.5.6. P.E.M.P. benefited from individual comments and discussions with Andrea Afruni and Fernanda Román-Oliveira. L.P. acknowledges support from the Centre National d’Études Spatiales (CNES) and from the European Research Council (ERC) under the European Unions Horizon 2020 research and innovation program (grant agreement No. 834148). G.P. acknowledges support from NOVA, Project 10.1.5.18. E.A.K.A. is supported by the WISE research programme, which is financed by the Netherlands Organization for Scientific Research (NWO). We used SIMBAD and ADS services, as well the Python packages NumPy (Oliphant 2007), Matplotlib (Hunter 2007), and SciPy (Virtanen et al. 2020), for which we are thankful.

References

- Bacchini, C., Fraternali, F., Iorio, G., & Pezzulli, G. 2019, *A&A*, **622**, A64
 Bacchini, C., Fraternali, F., Iorio, G., et al. 2020, *A&A*, **641**, A70
 Burstein, D., Bender, R., Faber, S., & Nolthenius, R. 1997, *AJ*, **114**, 1365
 Cappellari, M., Scott, N., Alatalo, K., et al. 2013, *MNRAS*, **432**, 1709
 Catinella, B., Saintonge, A., Janowiecki, S., et al. 2018, *MNRAS*, **476**, 875
 Cortese, L., Fogarty, L. M. R., Bekki, K., et al. 2016, *MNRAS*, **463**, 170
 Dalcanton, J. J., Spergel, D. N., & Summers, F. J. 1997, *ApJ*, **482**, 659
 di Teodoro, E. M., Posti, L., & Ogle, P. M. 2021, *MNRAS*, submitted
 Dutton, A. A., & van den Bosch, F. C. 2012, *MNRAS*, **421**, 608
 El-Badry, K., Quataert, E., Wetzel, A., et al. 2018, *MNRAS*, **473**, 1930
 Elson, E. C. 2017, *MNRAS*, **472**, 4551
 Fall, S. M. 1983, in *Internal Kinematics and Dynamics of Galaxies*, ed. E. Athanassoula, *IAU Symp.*, **100**, 391
 Fall, S. M., & Efstathiou, G. 1980, *MNRAS*, **193**, 189
 Fall, S. M., & Romanowsky, A. J. 2018, *ApJ*, **868**, 133
 Genel, S., Fall, S. M., Hernquist, L., et al. 2015, *ApJ*, **804**, L40
 Hunter, J. D. 2007, *Comput. Sci. Eng.*, **9**, 90
 Irodoutou, D., Thomas, P. A., Henriques, B. M., Sargent, M. T., & Hislop, J. M. 2019, *MNRAS*, **489**, 3609
 Jadhav, Y. V., & Banerjee, A. 2019, *MNRAS*, **488**, 547
 Krumholz, M. R., Leroy, A. K., & McKee, C. F. 2011, *ApJ*, **731**, 25
 Kurapati, S., Chengalur, J. N., Pustilnik, S., & Kamphuis, P. 2018, *MNRAS*, **479**, 228
 Lagos, C. d. P., Theuns, T., Stevens, A. R. H., et al. 2017, *MNRAS*, **464**, 3850
 Lelli, F., Fraternali, F., & Sancisi, R. 2010, *A&A*, **516**, A11
 Lelli, F., McGaugh, S. S., & Schombert, J. M. 2016, *AJ*, **152**, 157
 Leroy, A. K., Walter, F., Brinks, E., et al. 2008, *AJ*, **136**, 2782
 Lutz, K. A., Kilborn, V. A., Koribalski, B. S., et al. 2018, *MNRAS*, **476**, 3744
 Mancera Piña, P. E., Fraternali, F., Adams, E. A. K., et al. 2019, *ApJ*, **883**, L33
 Mancera Piña, P. E., Fraternali, F., Oman, K. A., et al. 2020, *MNRAS*, **495**, 3636
 Mancera Piña, P. E., Posti, L., Fraternali, F., Adams, E. A. K., & Oosterloo, T. 2021, *A&A*, **647**, A76
 Marconi, A., & Hunt, L. K. 2003, *ApJ*, **589**, L21
 McGaugh, S. S., Schombert, J. M., Bothun, G. D., & de Blok, W. J. G. 2000, *ApJ*, **533**, L99
 Mo, H. J., Mao, S., & White, S. D. M. 1998, *MNRAS*, **295**, 319
 Murugesan, C., Kilborn, V., Jarrett, T., et al. 2020, *MNRAS*, **496**, 2516
 Obreja, A., Stinson, G. S., Dutton, A. A., et al. 2016, *MNRAS*, **459**, 467
 Obreschkow, D., & Glazebrook, K. 2014, *ApJ*, **784**, 26
 Obreschkow, D., Glazebrook, K., Kilborn, V., & Lutz, K. 2016, *ApJ*, **824**, L26
 Ogle, P. M., Jarrett, T., Lanz, L., et al. 2019, *ApJ*, **884**, L11
 Oliphant, T. E. 2007, *Comput. Sci. Eng.*, **9**, 10
 Pedrosa, S. E., & Tissera, P. B. 2015, *A&A*, **584**, A43
 Ponomareva, A. A., Verheijen, M. A. W., & Bosma, A. 2016, *MNRAS*, **463**, 4052
 Ponomareva, A. A., Verheijen, M. A. W., Peletier, R. F., & Bosma, A. 2017, *MNRAS*, **469**, 2387
 Posti, L., Fraternali, F., Di Teodoro, E. M., & Pezzulli, G. 2018a, *A&A*, **612**, L6
 Posti, L., Pezzulli, G., Fraternali, F., & Di Teodoro, E. M. 2018b, *MNRAS*, **475**, 232
 Posti, L., Famaey, B., Pezzulli, G., et al. 2020, *A&A*, **644**, A76
 Robotham, A. S. G., & Obreschkow, D. 2015, *PASA*, **32**, e033
 Romanowsky, A. J., & Fall, S. M. 2012, *ApJS*, **203**, 17
 Romeo, A. B. 2020, *MNRAS*, **491**, 4843
 Salo, H., Laurikainen, E., Laine, J., et al. 2015, *ApJS*, **219**, 4
 Shi, Y., Zhang, Z.-Y., Wang, J., et al. 2021, *ApJ*, **909**, 20
 Stevens, A. R. H., Lagos, C. d. P., Obreschkow, D., & Sinha, M. 2018, *MNRAS*, **481**, 5543
 Tremmel, M., Karcher, M., Governato, F., et al. 2017, *MNRAS*, **470**, 1121
 Tully, R. B., & Fisher, J. R. 1977, *A&A*, **500**, 105
 Virtanen, P., Gommers, R., Oliphant, T. E., et al. 2020, *Nat. Methods*, **17**, 261
 Wang, J., Koribalski, B. S., Serra, P., et al. 2016, *MNRAS*, **460**, 2143
 Zoldan, A., De Lucia, G., Xie, L., Fontanot, F., & Hirschmann, M. 2018, *MNRAS*, **481**, 1376

Appendix A: The role of molecular gas

Table A.1. Same as Table 1 but taking into account the presence of molecular gas.

	α	β	γ	σ_{\perp}
Stars	0.67 ± 0.03	0.55 ± 0.08	-3.57 ± 0.23	0.11 ± 0.01
Gas	0.75 ± 0.03	-0.50 ± 0.04	-4.46 ± 0.25	0.08 ± 0.01
Baryons	0.72 ± 0.02	0.42 ± 0.05	-4.23 ± 0.19	0.08 ± 0.01

In this Letter we have assumed $M_{\text{gas}} = 1.33M_{\text{HI}}$ and $j_{\text{gas}} = j_{\text{HI}}$, neglecting any contribution from molecular (H_2) and ionised gas. Here we show that observationally motivated corrections to account for the presence of H_2 do not change our results.

To account for M_{H_2} in our estimates of M_{gas} , we relied on the results from Catinella et al. (2018), who provide measurements of the ratio $M_{\text{H}_2}/M_{\text{HI}}$ as a function of M_* for a large sample of nearby galaxies. We fitted a linear relation to their binned measurements (see their Fig. 9 and Table 3), finding

$$\log(M_{\text{H}_2}/M_{\odot}) = 0.26 \log(M_*/M_{\odot}) + \log(M_{\text{HI}}/M_{\odot}) - 3.4. \quad (\text{A.1})$$

The scatter is large, and we adopted an uncertainty of 50% in M_{H_2} . With this, we redefined $M_{\text{gas}} = 1.33(M_{\text{HI}} + M_{\text{H}_2})$ and updated M_{bar} and f_{gas} accordingly. For massive discs, the correction to M_{gas} is about 0.1 dex, which is smaller than the typical uncertainty in M_{gas} (~ 0.13 dex); the correction is even smaller for the dwarfs. The change in f_{gas} is of the same order, also being negligible at low masses and changing by up to 0.1 dex at the high-mass end; this change is also of the order of the typical uncertainty in f_{gas} . Correspondingly, M_{bar} remains largely unchanged since the correction is smaller than 0.04 dex at all masses.

Including H_2 can also affect j_{gas} , as seen from the equation

$$j_{\text{gas}} = f_{\text{atm}} j_{\text{HI}} + (1 - f_{\text{atm}}) j_{\text{H}_2}, \quad (\text{A.2})$$

where j_{HI} and j_{H_2} are the specific angular momenta of the atomic and molecular gas components, respectively, and f_{atm} is the atomic-to-total gas ratio, $f_{\text{atm}} = M_{\text{HI}}/(M_{\text{HI}} + M_{\text{H}_2})$.

Obreschkow & Glazebrook (2014) provide measurements of j_{HI} , j_{H_2} , and j_{gas} for a sample of 16 spiral galaxies. In addition to this, we computed j_{HI} for four galaxies in our sample that have surface densities and CO rotation curves available from Bacchini et al. (2020). The typical ratio between j_{gas} and j_{HI} is 0.85, which translates into a shift of 0.07 dex. Thus, on average, including H_2 implies a correction to j_{HI} such that $\log(j_{\text{gas}}) = \log(j_{\text{HI}}) - 0.07$. The correction is of the same order as the average uncertainty in j_{gas} , 0.08 dex.

We again fitted the 3D relations taking into account the above corrections to M_{gas} , f_{gas} , j_{gas} , and M_{bar} . The new coefficients for stars, gas, and baryons are listed in Table A.1. As expected, they are fully consistent with those reported in Table 1 within the uncertainties. Therefore, we conclude that including H_2 does not have a significant effect on the derivation of the j - M - f_{gas} laws, and our results remain unchanged.

Appendix B: The j_{gas}/j_* ratio from our best fitting planes

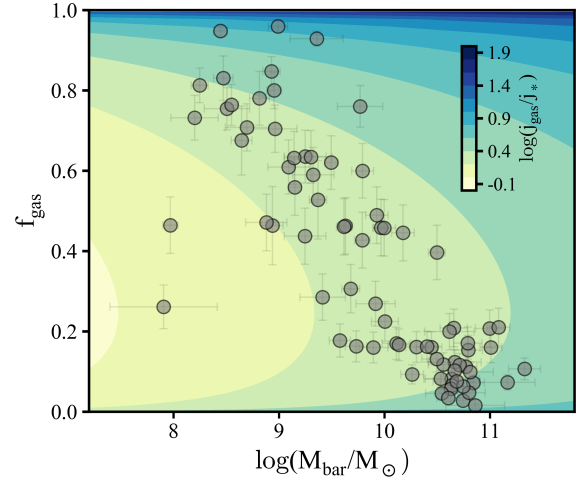


Fig. B.1. Relation between M_{bar} , f_{gas} , and the best fitting j_{gas}/j_* ratio according to Eq. (B.1). The background shows increasing levels of j_{gas}/j_* , and the grey points show our galaxies with convergent measurements of j_{gas} and j_* .

The j_{gas}/j_* ratio can be obtained directly from our individual galaxy measurements, as shown in Fig. 2. Nevertheless, j_{gas}/j_* can also be obtained by using our best fitting stellar and gas relations. This allows us to extrapolate j_{gas}/j_* to values of f_{gas} and M_{bar} beyond our observations and, in principle, to neglect the observational uncertainties since they are accounted for in our best fitting planes. Using Eq. (4) the ratio becomes

$$\log(j_{\text{gas}}/j_*) = \alpha_{\text{gas}} \log(f_{\text{gas}} M_{\text{bar}}) - \alpha_* \log[(1 - f_{\text{gas}}) M_{\text{bar}}] + (\beta_{\text{gas}} - \beta_*) \log(f_{\text{gas}}) + \gamma_{\text{gas}} - \gamma_*, \quad (\text{B.1})$$

and the corresponding surface according to Table 1 is shown in Fig. B.1. As can be seen, most of our galaxies lie within a region where $j_{\text{gas}}/j_* \sim 2$. It will be interesting to see where other large samples of galaxies would lie in Fig. B.1 and whether they also follow the expected dependence of j_{gas}/j_* on M_{bar} and f_{gas} .