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CATS: continuous turbulence characterization station for both optical link and astronomical support.

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ABSTRACT

The optical turbulence measurement and characterization has become an essential information for optical link (telecommunication, telemetry, time transfert, ...) and for high-angular resolution imaging in astronomy. For example, it has an impact on the quality and on the bit error rate of optical communication signals and it decrease the resolution of astronomical images. The J.L. Lagrange Laboratory of the Observatoire de la Côte d'Azur has a long and recognized expertise in atmospheric optics and turbulence for astronomical purposes. We have developed recently a new generation of autonomous instruments for daytime and nighttime atmospheric optical turbulence measurement. In 2015, a complete instrumental platform, the Calern Atmospheric Turbulence Station (CATS) has been installed at the Calern observatory. CATS is an automatic station equipped with a set of complementary instruments using original techniques for monitoring optical turbulence from the first meters above the ground up to the top of the atmosphere including the dome seeing. The station is based on 4 instruments. PML measures continuously the vertical profiles of the refractive index structure constant C_n^2 with a high vertical resolution ($\sim 100m$ at ground level). GDIMM monitors the wavefront coherence parameters (seeing, isoplanatic angle, coherence time, scintillation, wavefront coherence outer scale). A weather station provides the ground meteorological conditions (pressure, temperature, relative humidity, wind speed and direction). And the nighttime cloud fraction is given by an all-sky camera. Another instrument called INTENSE (INdoor TurbulENCE Sensor) is occasionally associated with CATS station to measure the seeing inside the dome of the 1.5m MeO telescope to evaluate its contribution to the whole turbulence. A new tool has been integrated to CATS station consisting in a forecasting model using the Weather Research and Forecasting (WRF) system coupled to a turbulence model to predict daytime and nighttime meteorological and optical turbulence conditions for the next 48h.

Keywords: Turbulence - Atmospheric effects - Instrumentation - Forecasting - Free space optical link

1. INTRODUCTION

Since decades the astronomical community is aware of the importance of the characterization of the atmosphere above observational site to optimize the location of instruments. In fact, atmospheric turbulence shakes the atmospheric layers of different refractive index. This phenomenon disturbs the light propagation through the atmosphere resulting in deformations and fluctuations of images acquired with telescopes.

More recently, with the development of the free space optical links, it has become apparent that turbulence also has a strong impact on signal quality, bit error rate and link budget.

For both astronomical and free space optical links, the adaptative optics is able to compensate part of the turbulence degradation. However, the dimensioning, conception and use depends on the turbulence characterization.

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Regarding, all these points, it is clear that having an instrumentation able to measure 24h a day the optical turbulence conditions is of a real interest for both astronomical and optical telecommunications communities.

In this goal, since 2015, the Observatoire de la Côte d’Azur (OCA) has built and installed a new generation station for atmospheric turbulence characterization on the Calern Observatory. The Calern Atmospheric Turbulence Station¹⁻³ (CATS) came out from a long and recognized expertise of the OCA/Lagrange Laboratory (OCA) in Atmospheric Optics. Indeed this laboratory played an important role in selection of the major astronomical sites over the world: GranTeCan (Canary Islands), Very Large Telescope (Chile), Southern Gemini (Chile), Keck, Northern Gemini & Subaru (Hawaii, USA). More recently, the Lagrange laboratory was involved in the European Extremely Large Telescope (E-ELT) and the Thirty Meter Telescope (TMT) site selections. Moreover, our team was also in charge of the qualification of the site of Dome C in Antarctica whose potential in Astronomy is considerable. This later was really challenging because of the peculiar conditions of this site.

The CATS station benefits of a set of a new generation of autonomous instruments allowing the measurements of optical turbulence during both daytime and nighttime. It is a support of other project on the Calern Observatory, mainly to improve the link budget of Laser Telemetry from MeO station and generally free space optical links. CATS is also supporting projects on MéO (Métrologie Optique) station and C2PU (Centre Pédagogique Planète Univers) telescopes to test and validate new concepts and components in order to overcome the current limitations of existing Adaptive Optics (AO) systems. The idea is to offer to AO community a fully operational on-sky test platform. The ultimate goal being the optimization of the scientific returns of the AO assisted instrumentations.

To go further, we have recently added into the CATS station a predictive model⁴⁻⁷ allowing to forecast 48h in advance the atmospheric and optical conditions above the Calern Observatory. This prediction tool will help, observers to plan their activities. This tool, fully automatic, will also be useful for the free space optical link because it will bring informations about the availability of the ground station to receive optical signals.

2. THE CALERN ATMOSPHERIC TURBULENCE STATION: CATS

The CATS station¹⁻³ (see figure 1) is a fully autonomous station working without human needs. It is composed of a set of complementary instruments allowing the measurements of both optical and meteorological parameters:

- A weather station measures the ground parameters such as the temperature, the relative humidity or the wind speed and direction, and gives informations for the aperture of the domes.
- An all-sky camera^{2,3} to measure the cloud cover, and command, with the weather station, the aperture of the domes.
- A Generalized DIMM (G-DIMM)⁸⁻¹⁰ (figure 1, left) measures the nighttime integrated optical parameters such as the seeing, the isoplanatic angle and the coherence time.
- A Profiler of Moon Limb (PML)¹¹⁻¹⁴ (figure 1, right) measures during daytime (on the sun limb) and nighttime (on the moon limb) the vertical profile of the C_n^2 with a high resolution ($\Delta h = 100m$ near the ground) and the integrated parameters (seeing, isoplanatic angle, wavefront coherence outer scale, scintillation, and, in the future, the coherence time)
- An INdoor TurbulENce SEnsor^{15,16} (INTENSE) can be sometimes used to estimate the turbulence inside the MéO station dome. This indoor turbulence can have a significant contribution to the degradation of the optical resolution of the telescope.

The main advantage of this station is that it acquires measurements during daytime and nighttime in an automatic way thanks to the weather station and the all-sky camera which control and allow the beginning of observations in function of the meteorological conditions. Therefore, since 2015 that CATS station has acquired a large database which has been used for many studies, and which will be used for ongoing and future works in both astronomy and optical telecommunications fields.



Figure 1. CATS station installed on the Calern Observatory. The left instrument is the G-DIMM, and the right one is the PML.

2.1 Generalized DIMM: GDIMM

The Differential Image Motion Monitor¹⁷ (DIMM) is a well known and widely used instrument because of its simplicity. It is based on a small telescope with an entrance pupil made of 2 small sub-apertures (6cm diameter) observing a bright single star with a short exposure time. The Generalized DIMM is a DIMM updated by adding a third sub-aperture at the entrance pupil (see Fig. 2). This third sub-aperture (10cm diameter and 4cm central obstruction) allows the instrument to deduce the isoplanatic angle and the wavefront coherence outer scale.



Figure 2. GDIMM instrument installed on the Calern Observatory.

The most important parameter measured by the GDIMM is the seeing, which is deduced from the Angle of Arrival (AA) fluctuations of the perturbed wavefront. These AA fluctuations are measured through the two 6cm sub-apertures with the well known DIMM method.

Another important parameter for astronomy and free space optical link is the isoplanatic angle θ_0 because it gives information on the angular distance from which the incoming wavefront is uncorrelated. It is deduced from the scintillation of a single star measured through a 10cm pupil with a 4 cm central obstruction. The principle of this θ_0 estimation method is based on the similarity of the theoretical expressions of θ_0 and the scintillation index s .^{18,19} However, Ref. 20 shown that the scintillation can be affected by the presence of the Moon not far from the observed bright star. Therefore, in such conditions, the isoplanatic angle deduced form scintillation could be also affected by the Moon light.

The most difficult parameter to measure is the wavefront coherence outer scale \mathcal{L}_0 . The GDIMM instrument uses the variances of AA fluctuations through different size sub-apertures by means of the von Karman model (see Ref. 3 for more details).

In the other hand, the coherence time τ_0 relevant for adaptive optics and interferometry, is given by $\tau_0 = 0.31 \frac{r_0}{\bar{v}}$, where \bar{v} the effective wind speed, is a weighted average of the wind speed on the whole atmosphere. r_0 is the well known Frieds parameter. Refs. 3, 10 have shown that it is possible to deduce the effective wind speed from the temporal structure functions of the AA fluctuations.

The GDIMM instrument is fully automatic and all parameters are measured and computed in real time and are uploaded on the CATS website* (see Fig. 3). This website is daily used by observers in Calern Observatory in parallel of their observations.

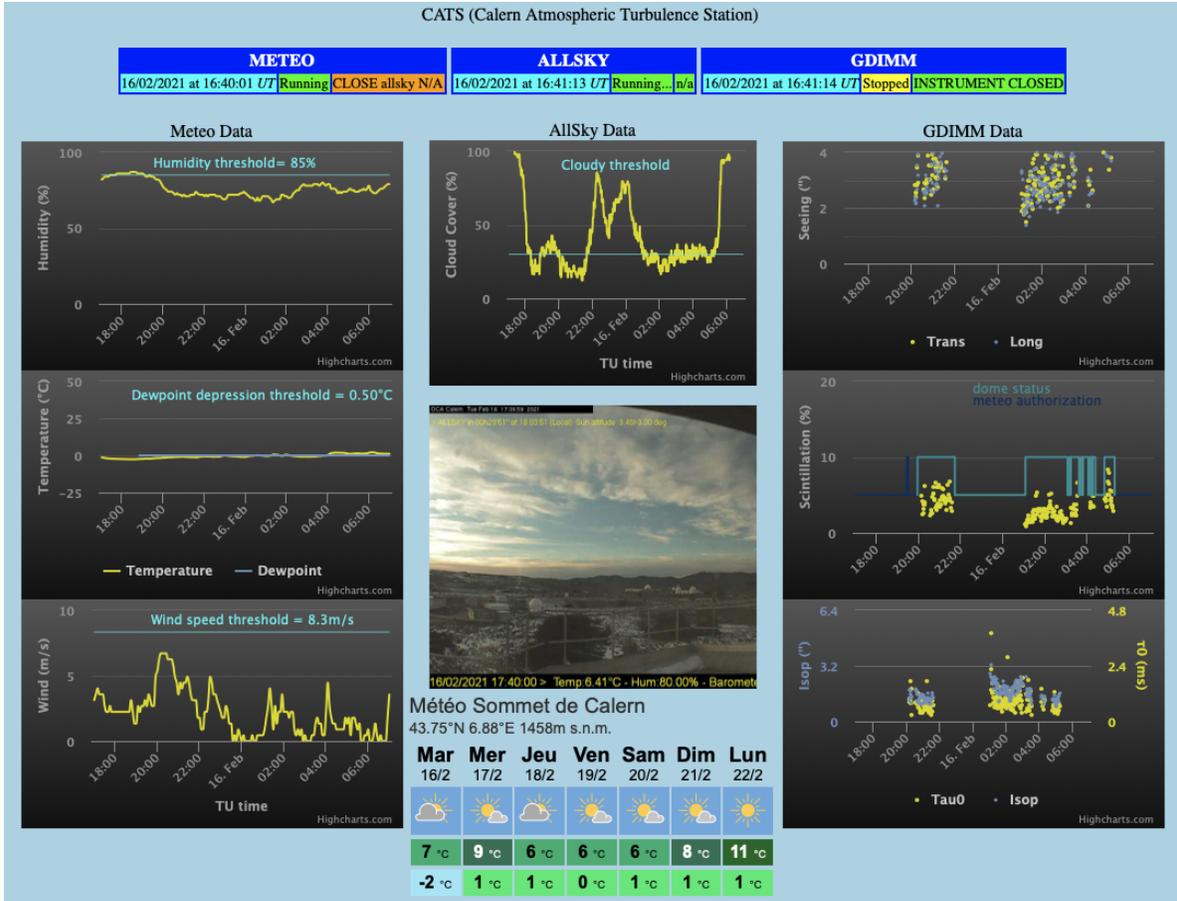


Figure 3. GDIMM results uploaded in real time on the CATS website.

2.2 The Profiler of Moon Limb: PML

The PML instrument is a profiler of the C_n^2 with the high vertical resolution. The C_n^2 is the structure constant of the refractive index fluctuations of the atmosphere. It gives information of the turbulence intensity of each layer. The PML instrument is based on a differential method by observation of the lunar or solar limb through two sub-apertures (see Fig. 4). The Moon or Solar limb acts as a continuum of double stars with all possible angular separations required between two points to scan the atmosphere with a very fine resolution. Doing measurements on both Sun and Moon, the PML instrument gives also access to daytime measurements which are really important for free space optical links.

*<https://cats.oca.eu>

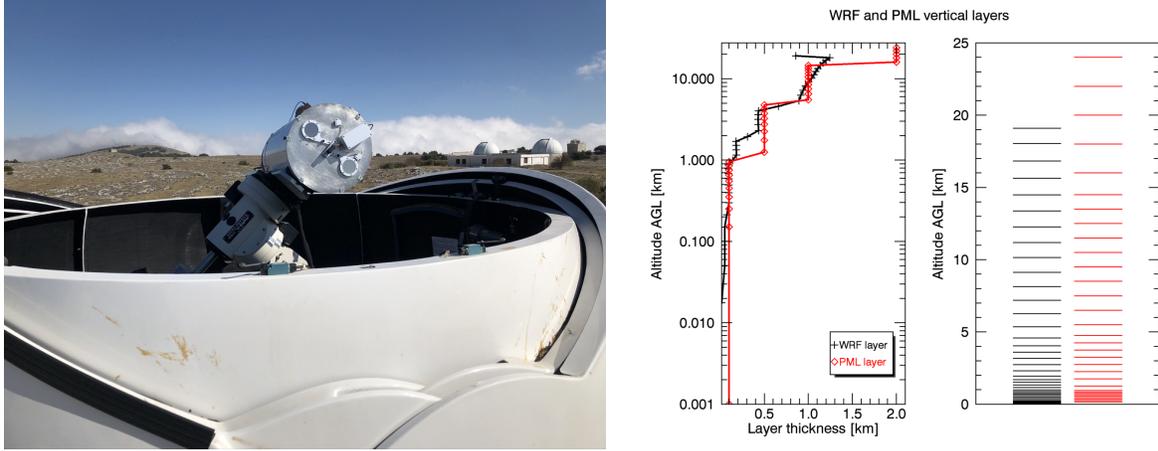


Figure 4. Left: PML instrument installed on the Calern Observatory. Right: Vertical resolution reached by the PML (red) and by the forecasting tool (black) (see section 2.3)

The vertical profile of the C_n^2 is deduced from the measurement of the angular correlation of the fluctuation differences in the wavefront AA deduced from the motion of the Moon (or Sun) limb image. The AA fluctuations are measured perpendicularly to the lunar (or solar) limb leading to transverse correlations for different angular separations along the Moon (or Sun). For more detail about the computation method, see Refs. 3, 14. Fig. 4, right, shows that the highest vertical resolution, $\Delta h = 100m$, is reached for the ground layer and it decreases progressively to $\Delta h = 2000m$ for the highest layers ($h > 15km$).

From this vertical profile, it is easy to compute integrated parameters such as the seeing or the isoplanatic angle. This later is deduced directly from the C_n^2 profile and it is not biased like the one deduced by the GDIMM with the presence of the Moon (see section 2.1).

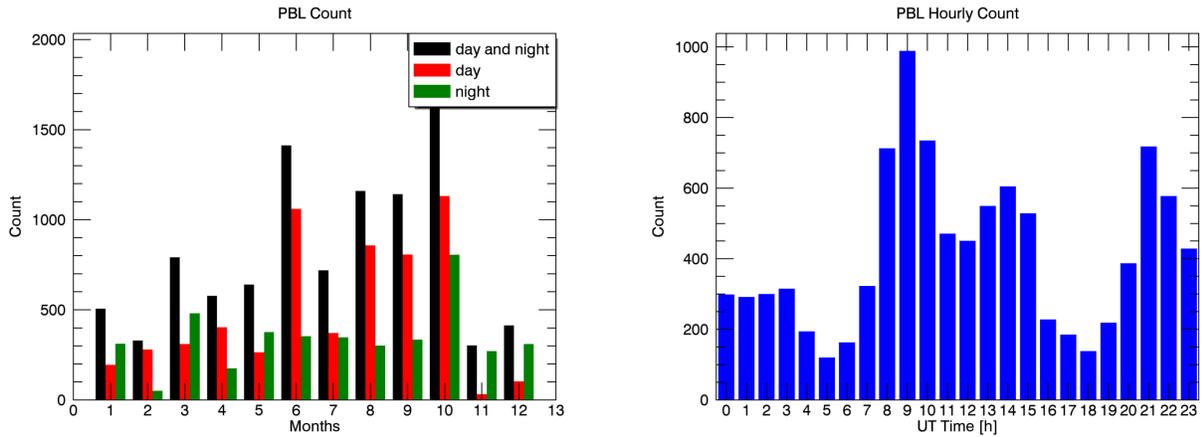


Figure 5. PML data distribution between 2015 and 2019. Left: monthly count. Right: Hourly count.

Since 2015, the PML instrument has acquired a large database which has been used within several astronomical and optical telecommunication projects. The figure 5 shows the PML data distribution computed monthly and hourly between 2015 and 2019. It appears that there are more measurements during the summer and during the day. This is mainly due to the fact that the Moon is not always visible.

2.3 Calern forecasting

The last tool we have decided to implement in the CATS station is the forecasting of atmospheric and turbulence conditions. We use the non-hydrostatic mesoscale model called Weather Research and Forecasting (WRF) developed in the National Center for Atmospheric Research (NCAR), USA.²¹ This model allows us to predict weather conditions within a tridimensional domain. From these weather forecasts we use an empirical model deduced from the analysis of balloon radio-sounding launched around the world.^{4-7,22} This model called BDTM (Balloon Derived Turbulence Model) follows equations hereafter:

$$C_T^2(h) = \phi(h)\chi(h)s(h)^{1/2} \quad (1)$$

$$\chi(h) = \frac{\delta\theta}{\delta h} \quad (2)$$

$$s(h) = \left(\frac{dU^2}{dh} + \frac{dV^2}{dh} \right)^{1/2} \quad (3)$$

where the C_T^2 is the structure constant of temperature fluctuations, χ is the vertical gradient of the potential temperature θ , h is the altitude, U and V are the wind speed respectively in the direction West-East and South-North, and s is the horizontal wind shear. $\phi(h)$ is the vertical profile deduced empirically from the analysis of the balloons radio-sounding,

$$\phi(h) = \frac{\langle C_T^2(h) \rangle_m}{\langle \chi(h) \rangle_m \langle s(h) \rangle_m^{1/2}} \quad (4)$$

where $\langle \rangle_m$ is the median value.

Finally, from C_T^2 we can deduce the C_n^2 following the well-known Gladstone's formula and others integrated parameters (seeing, isoplanatic angle, coherence time).

$$C_n^2(h) = \left(\frac{80.10^{-6} P(h)}{T(h)^2} \right)^2 C_T^2(h) \quad (5)$$

where $P(h)$ and $T(h)$ are respectively the vertical profiles of the atmospheric pressure and absolute temperature.

Since the end of July 2019, we have implemented a fully automatic prediction tool above the Calern Observatory using the aforementioned models. This prediction covers the next 48h and uploads the results on the CATS website (see Fig. 6).

In Ref. 7, we have used these predictions and the CATS database to constrain the empirical model by local measurements. This upgrade, called SL (Site Learning) method, brings a real improvement as shown in Fig. 7. Indeed, regarding these plots it appears that the predicted seeing is closer to the measured one and the dispersion is lower. This SL method is continuously improved thanks to the continuous expansion of the database.

The last ongoing implementation about the predictions concerns the possibility to use our large database in Machine Learning (ML) algorithms. This possibility offers the advantage to have short-term predictions (<2h) with a better accuracy than the one obtained from forecasting models. For more details about this tool currently in development, see Ref. 23.

In parallel, an ongoing study tries to retrieve the vertical profiles of the outer scale L_0 from measurements and predictions of weather conditions. This work could allow us to use also a theoretical model in complement of the empirical method presented here above. A. Rafalimanana will show the first results of this study in this ICSO conference and within the corresponding proceedings.

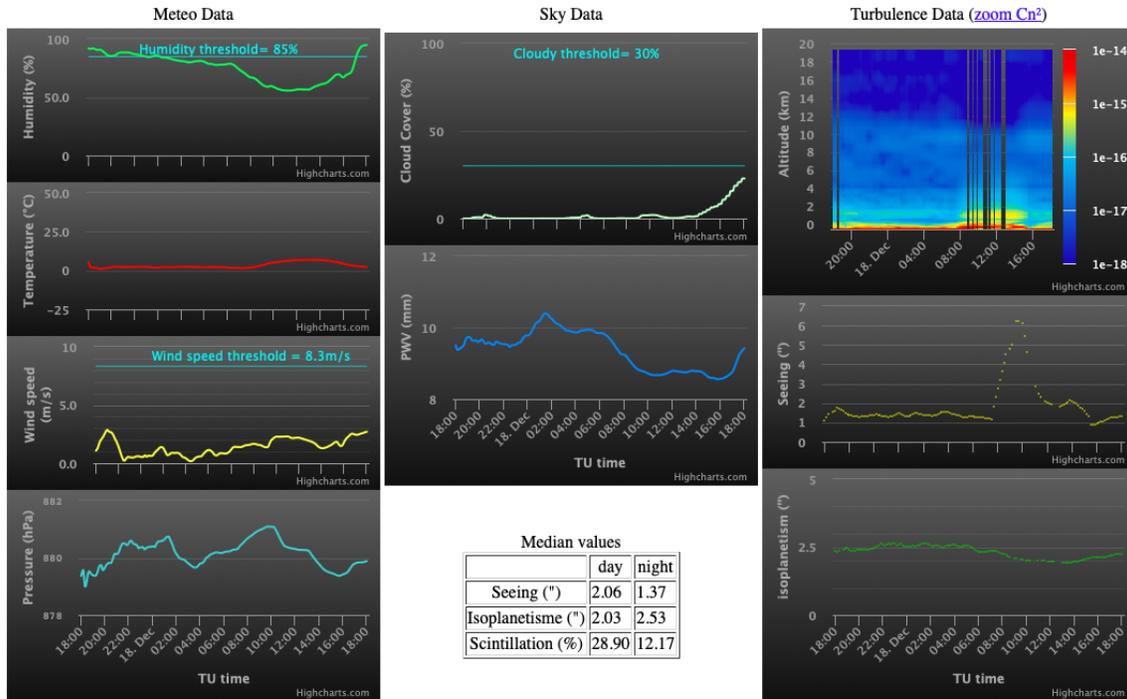


Figure 6. Calern Forecasting results uploaded every day on the CATS website.

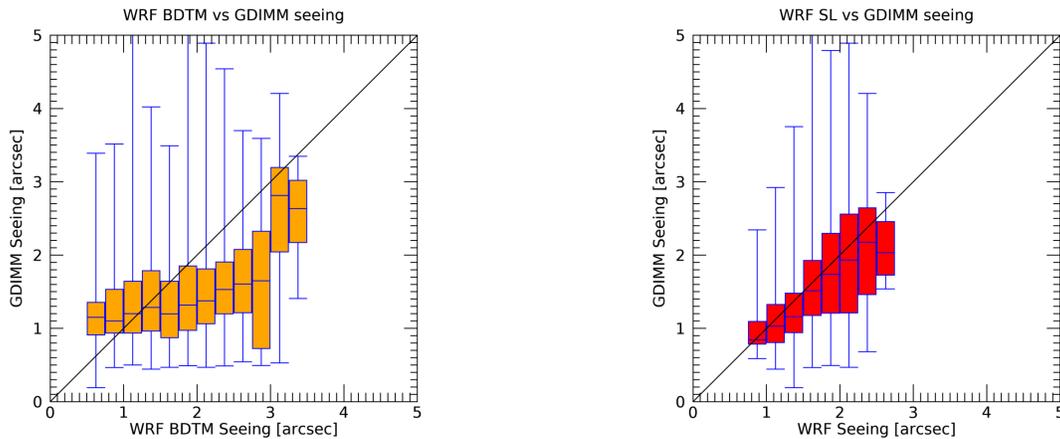


Figure 7. Scatter plot shown as box plot of the seeing forecasted using both BDTM (left) and SL (right) method and measured with GDIMM. The box corresponds to interval between first and third quartile. Within the box, the median values is plotted. The whiskers corresponds to extrema.

3. DISCUSSION

The CATS station is installed and operational at the Calern Observatory since autumn 2015. CATS is fully autonomous and monitors atmospheric turbulence parameters during both daytime and nighttime thanks to a set of new generation instruments. In particular CATS is equipped with the PML which is a profiler of turbulence measuring the conditions from the first meters above the ground to the borders of the atmosphere with the highest resolution known and during daytime and nighttime.

In addition, since more than 18 months, CATS is also equipped with a forecasting tool doing daily predictions of all atmospheric conditions (weather and turbulence) for the next 48h. These predictions have been tested and optimized thanks to a site learning method 7. In addition, a new tool using both the CATS database and the aforementioned 48h predictions is in development. This later uses machine learning algorithm to compute a short-terms prediction in order to increase the accuracy of the mid-terms ones²³. These predictions are of a real interest for observers at Calern because they give them an important information in advance to allow them to plan their observations. Moreover, such a prediction tool is important for free space optical link context because it will give information in advance on the availability of the ground station to receive signals from satellites.

All these points are the reason why the CATS station is involved in many scientific and educational projects at the Calern Observatory such as:

- study of turbulence impact on the laser links of the MéO Laser Ranging Station;
- development of Adaptive Optics projects;
- study of laser uplink propagation in the atmosphere thanks to the T2L2 space mission;
- campaign involving optical telecommunication from space to ground with the SOTA satellite and MéO laser ranging station.
- an ongoing European project with ESA uses the CATS station as a reference station in an optical telecommunication context.

Another campaign has been performed with the Optical PAYload for Lasercomm Science (OPALS) on International Space Station in spring 2016 and recently with the DLR OSIRIS instrument at the end of 2019. For these types of campaign, the PML instrument is very useful to understand the link budget of the lunar laser ranging activity at the MéO telescope. In addition, several other projects regarding adaptive optics systems are planned at the MéO telescope and C2PU telescope and they will benefit of the data given by the CATS station.

Finally, the CATS station is used each year for on sky practical work for students of the Université Côte d’Azur (UCA).

4. ACKNOWLEDGMENTS

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