



HAL
open science

Star formation triggered by galaxy interactions in modified gravity

Florent Renaud, Benoit Famaey, Pavel Kroupa

► **To cite this version:**

Florent Renaud, Benoit Famaey, Pavel Kroupa. Star formation triggered by galaxy interactions in modified gravity. *Monthly Notices of the Royal Astronomical Society*, 2016, 463 (4), pp.3637-3652. 10.1093/mnras/stw2331 . hal-03146786

HAL Id: hal-03146786

<https://hal.science/hal-03146786>

Submitted on 19 Feb 2021

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Star formation triggered by galaxy interactions in modified gravity

Florent Renaud,¹★ Benoit Famaey² and Pavel Kroupa^{3,4}

¹*Department of Physics, University of Surrey, Guildford GU2 7XH, UK*

²*Observatoire Astronomique de Strasbourg, Université de Strasbourg, CNRS UMR 7550, 11 rue de l'Université, F-67000 Strasbourg, France*

³*Helmholtz-Institut für Strahlen- und Kernphysik, Nussallee 14-16, D-53115 Bonn, Germany*

⁴*Astronomical Institute, Faculty of Mathematics and Physics, Charles University in Prague, V Holešovičkách 2, CZ-180 00 Praha 8, Czech Republic*

Accepted 2016 September 12. Received 2016 September 12; in original form 2016 July 24

ABSTRACT

Together with interstellar turbulence, gravitation is one key player in star formation. It acts both at galactic scales in the assembly of gas into dense clouds and inside those structures for their collapse and the formation of pre-stellar cores. To understand to what extent the large-scale dynamics govern the star formation activity of galaxies, we present hydrodynamical simulations in which we generalize the behaviour of gravity to make it differ from Newtonian dynamics in the low-acceleration regime. We focus on the extreme cases of interacting galaxies, and compare the evolution of galaxy pairs in the dark matter paradigm to that in the Milgromian dynamics (MOND) framework. Following up on the seminal work by Tiret & Combes, this paper documents the first simulations of galaxy encounters in MOND with a detailed Eulerian hydrodynamical treatment of baryonic physics, including star formation and stellar feedback. We show that similar morphologies of the interacting systems can be produced by both the dark matter and MOND formalisms, but require a much slower orbital velocity in the MOND case. Furthermore, we find that the star formation activity and history are significantly more extended in space and time in MOND interactions, in particular in the tidal debris. Such differences could be used as observational diagnostics and make interacting galaxies prime objects in the study of the nature of gravitation at galactic scales.

Key words: gravitation – methods: numerical – stars: formation – galaxies: interactions – galaxies: starburst.

1 INTRODUCTION

Interactions mark milestones in the evolution of galaxies by modifying their mass, stellar, gaseous and chemical contents, morphology, kinematics and dynamical properties (see e.g. Arp 1966; Sanders & Mirabel 1996; Springel & White 1999; Struck 1999; Saintonge et al. 2012; Duc & Renaud 2013, among many others). These events are often (but not always) associated with burst(s) of star formation such that, in the local Universe, all the most luminous galaxies (e.g. $> 10^{12} L_{\odot}$ for the ultraluminous infrared galaxies, ULIRGs; Houck et al. 1985; Kennicutt 1998) yield the signatures of major interactions (Armus, Heckman & Miley 1987; Ellison et al. 2013). Numbers of studies in all wavebands have characterized the properties of interacting systems, in particular their star formation activity, with the aim of pinning down the underlying physical processes (e.g. Schombert, Wallin & Struck-Marcell 1990; Hibbard & Mihos 1995; Bournaud et al. 2004; Chien et al. 2007; Smith et al. 2010; Boquien et al. 2011; Saintonge et al. 2012; Scudder et al. 2012).

In these fast evolving objects with complex geometries, numerical simulations have long been considered as a fundamental complement to observations. Starting with Toomre & Toomre (1972), all works point out the paramount role of gravitation in affecting both the galactic-scale structures (e.g. Barnes & Hernquist 1992; Quinn, Hernquist & Fullagar 1993; Dubinski, Mihos & Hernquist 1996; Mihos, Dubinski & Hernquist 1998; Di Matteo et al. 2007; Hopkins et al. 2009; Renaud et al. 2009; Moreno et al. 2013; Privon et al. 2013) and the internal, small-scale, physics of interacting galaxies (e.g. Barnes & Hernquist 1991; Chien & Barnes 2010; Teyssier, Chapon & Bournaud 2010; Hopkins 2013; Renaud et al. 2014). The scale-free aspect of gravitation makes it indeed a key process at galactic scale in the shaping of galaxies and their structures (spiral, bars, large-scale flows), at sub-galactic scales in the assembly of molecular clouds (Mac Low & Klessen 2004; Hennebelle & Chabrier 2008) and down to the scale of pre-stellar cores in star formation (Bate & Bonnell 2005; Bonnell, Dobbs & Smith 2013). Galaxy interactions are thus the perfect benchmark to understand the role of gravitation in star formation, at both galactic and sub-galactic scales.

The classical framework in which theoretical galactic studies are performed these days is the Λ cold dark matter (Λ CDM) paradigm.

* E-mail: florent.renaud@gmail.com

However, both the cosmological constant Λ and the CDM part of the model could also be related to a modification of gravity. On galaxy scales, the model is indeed plagued by severe problems, the most famous ones being the cusp–core problem (de Blok 2010; Oman et al. 2015, but see also Read et al. 2016), the too-big-to-fail problem (Boylan-Kolchin, Bullock & Kaplinghat 2011; Papastergis et al. 2015; Pawlowski et al. 2015) or the satellite planes problem (Kroupa, Theis & Boily 2005; Metz, Kroupa & Jerjen 2007; Metz, Kroupa & Libeskind 2008; Pawlowski, Pflamm-Altenburg & Kroupa 2012; Ibata et al. 2013, 2014; Pawlowski et al. 2015). There is also a more general problem linked to the finely tuned relation between the distribution of baryons and the gravitational field in galaxies, as encapsulated in various scaling relations involving a universal acceleration constant $a_0 \approx 10^{-10} \text{ m s}^{-2}$, including the tight baryonic Tully–Fisher relation (McGaugh et al. 2000; Lelli, McGaugh & Schombert 2016a; Papastergis, Adams & van der Hulst 2016), the diversity of shapes of rotation curves at a given maximum velocity scale (Oman et al. 2015) or the relation between the stellar and dynamical surface densities in the central regions of galaxies (Lelli et al. 2016b; Milgrom 2016), and many others (Famaey & McGaugh 2012). All this points to things happening as if the effects usually attributed to CDM on galaxy scales were actually due to a modified force law. The a priori simplest explanation for this would be that gravity is indeed effectively different in the weak field regime and accounts for the effects usually attributed to CDM. This paradigm is known as modified Newtonian dynamics (MOND), or Milgromian dynamics, suggested more than 30 years ago by Milgrom (1983). It predicted all the observed galaxy scaling relations well before they were precisely assessed by observations (Famaey & McGaugh 2012). Nevertheless, this paradigm cannot be complete, as a full theory of gravitation, also valid on cosmological scales, has not yet been found. But while successful on galaxy scales, the MOND paradigm has still been far from being fully explored even on these scales where it is currently successful. Hence, there is still potential for falsification of this paradigm in its a priori domain of validity. The main reason for this lack of exploration of all predictions of MOND on galaxy scales is its non-linear nature, and the previous lack of numerical codes devised to model galaxies in this framework.

After the pioneering work of Brada & Milgrom (1999), only a handful of codes have been designed (Nipoti, Londrillo & Ciotti 2007; Tiret & Combes 2007; Linares, Knebe & Zhao 2008; Angus et al. 2012), but all with their own caveats, notably regarding the treatment of hydrodynamics. The first treatment of gas in MOND simulations was proposed by Tiret & Combes (2008b), who used sticky particles, but a full hydrodynamical treatment of gas has become possible only recently thanks to the patches of the RAMSES code developed by Lügghausen, Famaey & Kroupa (2015) and Candlish, Smith & Fellhauer (2015). This is particularly important in order to study star formation in general, and in galaxy interactions in particular. The role of gravitation in star formation is central at both galactic and sub-galactic scales, and modifying it must have potentially observational consequences.

For instance, during interactions, tides can induce the formation of tails expanding up to ~ 100 kpc away from their progenitor galaxies. Some regions of these tails can become unstable, fragment and form stellar objects as massive as dwarf galaxies ($\sim 10^{8-9} M_\odot$). However, self-gravity is usually too weak to assemble such tidal dwarf galaxies (TDGs), and an external contribution to the local potential well is required. With Λ CDM, only the potential well of the galactic dark matter (DM) halo can have such catalyst effect and thus, it must be sufficiently extended to embed the TDGs along

the long tidal tails (Bournaud, Duc & Masset 2003). But since the tidal debris originates from the discs (and thus contains very little DM, if any) and because the surrounding DM haloes are dynamically hot, the DM distribution does not follow the baryonic one, and their external effect on TDG seeds remains mild. In MOND however, the baryonic seeds of TDGs generate their own ‘phantom DM’ and thus, an additional potential well. Therefore, instabilities leading to the formation of TDG seeds are strengthened by the local MOND potential which amplifies them and allows them to grow. As a consequence, the formation of TDGs is eased in the MOND framework, compared to the Newton case (Tiret & Combes 2008a; Combes & Tiret 2010). Since it provides a test of the gravitation paradigm, the nature of observed TDG candidates is intensively debated, in particular in the context of the Tully–Fisher relation (see e.g. Gentile et al. 2007; Lelli et al. 2015; Flores et al. 2016) and regarding the potential origin of satellite galaxies sitting on satellite planes in the context of MOND, which could actually be old TDGs instead of primordially formed dwarf galaxies (e.g. Kroupa et al. 2010; Kroupa 2015).

In this paper, we thus use the MOND framework to analyse the role of gravitation in enhancing star formation in interacting galaxies. We characterize the starburst activity associated with interactions in the context of MOND, and compare to that obtained when they are surrounded by haloes of particle DM. We present our suite of simulations in Section 2 and the results in Section 3. Conclusions are drawn in Section 4.

2 FORMALISM AND NUMERICAL METHOD

2.1 Modified gravity

In highly symmetrical situations (such as spherical symmetry), the net MOND gravitational acceleration \mathbf{g} is connected to the Newtonian term \mathbf{g}_N through

$$\mathbf{g} = \nu \left(\frac{g_N}{a_0} \right) \mathbf{g}_N, \quad (1)$$

where a_0 is a constant acceleration, generally chosen to be $a_0 \approx 10^{-10} \text{ m s}^{-2}$, and ν is an interpolation function such that

$$\begin{cases} \nu(x) \rightarrow 1 & \text{for } x \gg 1 \text{ (Newtonian regime)} \\ \nu(x) \rightarrow x^{-1/2} & \text{for } x \ll 1 \text{ (MOND regime)} \end{cases}. \quad (2)$$

In the weak field (or MOND) regime, one gets

$$\mathbf{g} \rightarrow \sqrt{a_0 \mathbf{g}_N}, \quad (3)$$

which successfully predicts the observed rotation curves of galaxies (see Famaey & McGaugh 2012, for a review). This relation indeed predicts for instance that galaxies of the same baryonic mass would share the same asymptotic circular velocity in accordance with the baryonic Tully–Fisher relation, but also that the central slope of the rotation curve of disc galaxies scales as $(dV/dR)_0 \propto \sqrt{\nu \rho_b(0)}$, where V is the circular velocity, R is the galactocentric distance, $\rho_b(0)$ is the central baryonic density (which is typically the central baryonic surface density over twice the scaleheight) and ν is taken close to the centre (at the first measured point of the rotation curve). Since ν depends on g_N , it also scales with the baryonic surface density, and a prediction of MOND is consequently that the central circular velocity slope depends mainly on the central baryonic surface density of the galaxy. Hence, MOND naturally predicts the diversity of rotation curve shapes at a given mass scale noted by Oman et al. (2015), solely from the existence, at a given mass

scale, of a diversity of central baryonic surface densities, and it also predicts a strong correlation of the latter with the circular velocity gradient $(dV/dR)_0$, as observed by Lelli, Fraternali & Verheijen (2013).

Outside of spherical symmetry, the previous formula must nevertheless be modified. This is usually done by modifying the Poisson equation, following the least action principle for a modified Lagrangian of gravitation. One such flavour of MOND, called quasi-linear MOND (QUMOND; Milgrom 2010), provides a formalism similar to the Newtonian case by introducing a ‘phantom DM’ density which is fully determined by the baryonic distribution.

Following Famaey & Binney (2005), for galaxies,¹ we can choose an interpolation function²

$$v(x) = \frac{1}{2} + \frac{\sqrt{x^2 + 4x}}{2x}. \quad (4)$$

Then, following Famaey & McGaugh (2012), we define the function³

$$\tilde{v}(x) = v(x) - 1 \quad (5)$$

that allows us to write the modified Poisson equation

$$\nabla^2 \phi(\mathbf{x}) = 4\pi G \rho_b(\mathbf{x}) + \nabla \cdot \left[\tilde{v} \left(\frac{|\nabla \phi_N|}{a_0} \right) \nabla \phi_N(\mathbf{x}) \right], \quad (6)$$

where ϕ and ϕ_N are the net and Newtonian potentials, respectively, and ρ_b is the baryonic density, which fulfils the Newtonian Poisson equation $\nabla^2 \phi_N(\mathbf{x}) = 4\pi G \rho_b(\mathbf{x})$. We then introduce the ‘phantom DM’ density ρ_{ph}

$$\nabla^2 \phi(\mathbf{x}) = 4\pi G [\rho_b(\mathbf{x}) + \rho_{\text{ph}}(\mathbf{x})], \quad (7)$$

which is fully defined once the baryonic distribution is known, and can be seen as the (non-particle) MOND equivalent to the DM contribution in the classical case. In the rest of the paper, we adopt the QUMOND formalism and the standard value $a_0 = 1.12 \times 10^{-10} \text{ m s}^{-2}$.

Since MOND has been intensively used in the last decades to mainly study the dynamics of the outer regions of galaxies and galaxy clusters, it is often wrongly seen as a modification of Newtonian dynamics at large scales ($\sim 100 \text{ kpc} - 1 \text{ Mpc}$). It is however important to keep in mind that MOND does not directly depend on spatial scale but on the local gravitational acceleration. Therefore, it plays an important role where Newtonian gravitation is weak (parametrized by a_0), which is not strictly equivalent as playing an important role in the outskirts of galaxies, or in galaxy clusters, like DM does.

2.2 Numerical method

A handful of simulation codes have been developed within the MOND framework to address questions raised by Milgromian gravity (Brada & Milgrom 1999; Nipoti et al. 2007; Llinares et al. 2008; Tiret & Combes 2008a; Candlish et al. 2015; Lüghausen et al. 2015). However, in the era of multi-scale and multi-physics numerical studies, it is necessary to include the generalization of the gravitation law in a broader context, notably to correctly treat the hydrodynamics and the physics of star formation. That is why not

only Poisson solvers but also more versatile simulation codes originally designed in the Newtonian framework have recently been modified to solve the generalized Poisson’s equation of MOND. Most notably, Lüghausen et al. (2015) proposed a patch named PHANTOM OF RAMSES (POR) to the adaptive mesh refinement (AMR) RAMSES code (Teyssier 2002). No other functionalities or implementations of RAMSES than the Poisson solver are affected by the use of POR.

All the simulations presented here are performed using the AMR RAMSES code and the POR patch. Any other physical aspect is treated as in Renaud, Bournaud & Duc (2015a). In short, the galaxy systems (isolated or in pairs) are modelled without accretion of external gas nor DM. Heating comes from ultraviolet radiation of cosmic origin (tabulated at redshift zero as in the public version of RAMSES) and stellar feedback (see below), and the atomic cooling used is tabulated at solar metallicity (Courty & Alimi 2004). Gas denser than a density threshold of 0.6 cm^{-3} is converted into star particles, at a fixed efficiency per free-fall time (4 per cent; see also Renaud et al. 2013). The threshold combined with the efficiency corresponds to a star formation rate (SFR) of the order of $1 M_{\odot} \text{ yr}^{-1}$ for isolated galaxies, i.e. typical of the main-sequence galaxies in the local Universe (e.g. Daddi et al. 2010b). Star formation is only active in the densest regions of the galaxies, i.e. the most refined volumes of the simulations. Appendix A shows that in these regions, small-scale dynamics are dominated by the baryonic component and are fairly independent of the gravitation paradigm adopted. The self-gravity of such systems, responsible for cloud collapse and fragmentation leading to star formation, lies in the non-modified regime, or strong field, of MOND. Therefore, we use the exact same subgrid recipe for star formation for both the Newtonian and Milgromian runs. Differences in the star formation activities would thus originate from the formation of the star-forming clouds themselves due to larger scale dynamics like large-scale gas flows, compression, shocks, shear, etc.

Stellar feedback includes photoionization in H II regions, radiation pressure (Renaud et al. 2013) and thermal Type II supernovae (Dubois & Teyssier 2008; Teyssier et al. 2013). The simulated volume spans $(400 \text{ kpc})^3$, with Dirichlet’s boundary conditions, and with the highest resolution of the AMR grid being 6 pc. Additional tests show that our results are not affected by the choice of size for the simulation volume.

The simulations have been performed on the *Curie* supercomputer hosted at the *Très Grand Centre de Calcul* (TGCC).

2.3 Simulation suite

In the following, all galaxy models are the same (apart from the presence/absence of a DM halo), as described in Section 2.5. The differences in the simulation parameters are solely based on the orbits. We have performed four simulations of interacting galaxies, as follows.

(i) Using the orbit of Renaud et al. (2015a) designed to reproduce the morphology and kinematics of the Antennae galaxies (NGC 4038/39), we run a simulation in the Newtonian framework using the MW-A model (Kuijken & Dubinski 1995) setup for both progenitors as described in Section 2.5. Using this galaxy model introduces slight differences from Renaud et al. (2015a) in the final result, and the match to the Antennae system is, de facto, not as good as in previous works. This simulation is labelled N_A hereafter. The equivalent MOND simulation, labelled M_A , is set up using the exact same parameters: same orbit and same progenitors, only with the

¹ See however Hees et al. (2016) for constraints in the Solar system.

² This function is equivalent to the inverse interpolation function μ in equation 2 of Candlish et al. (2015).

³ Our \tilde{v} equals the function ν of Lüghausen et al. (2015).

Table 1. Initial orbital parameters. ‘*N*’ stands for Newton and ‘*M*’ for MOND. ‘*A*’ indicates the orbit of the Antennae model of Renaud et al. (2015a), and ‘*S*’ represents lower initial velocities (‘Slow’).

Simulation	N_A and M_A	N_S and M_S
<i>Galaxy 1</i>		
Position (kpc)	(12.7, −30.3, 46.7)	=
Velocity (km s ^{−1})	(−26.9, 23.3, −71.5)	×0.5
Spin axis	(0.67, −0.71, 0.20)	=
<i>Galaxy 2</i>		
Position (kpc)	(−12.7, 30.3, −46.7)	=
Velocity (km s ^{−1})	(26.9, −23.3, 71.5)	×0.5
Spin axis	(0.65, 0.65, −0.40)	=

DM particles removed and replaced by the QUMOND gravitation, as detailed in Sections 2.1 and 2.5. The system obtained is thus *not* tailored to reproduce any observation of the real Antennae. It however provides a comparison as direct as possible with the Newtonian case. For these simulations, the initial relative velocity of the progenitors is ≈ 160 km s^{−1}.

(ii) We design another MOND simulation, named M_S (for ‘slow’), providing a better fit to the Antennae, although precisely matching all features of the real system is not our goal. Keeping the progenitors unchanged (for simplicity), the main objective of this run is to compensate for the differences in dynamical friction between Newton and MOND (e.g. Kroupa 2015), to reduce the orbital period and to adjust the length of the tidal tails (see below). All these targets have been reached by decreasing the initial orbital velocity of the progenitors by a factor of 2 (i.e. down to ≈ 80 km s^{−1}).⁴ For the sake of completeness, we finally run the Newtonian equivalent (N_S) using the same initial parameters as in M_S .

The initial conditions of the interactions are listed in Table 1. The angles between the spin axis of the galaxies and the vector normal to the orbital plane are 46° and 58° for the progenitors 1 and 2, respectively, meaning that the interaction is prograde (i.e. with spin inclination <90°) for both galaxies, but with a spin–orbit coupling slightly stronger for galaxy 1 than for galaxy 2 (prograde encounters favour the formation of long tidal features; see e.g. Duc & Renaud 2013).

The progenitor galaxies are initially placed at a large distance from their encounter position (~ 60 kpc each). This ensures that, for the velocities adopted, (i) they start in a quasi-isolation stage (in agreement with the initial setup of the models), (ii) they have enough time to virialize and evacuate the imperfections of the initial conditions before the interaction itself, and (iii) they interact before the formation of substructures (bar, spirals) in the discs.

We also run Newton and MOND simulations with an initial velocity 1.75 times larger than N_A and M_A (not shown in this paper for the sake of clarity). These complementary simulations lead to the same conclusions as those presented below.

2.4 Comparing DM and MOND simulations

Ideally, to allow for direct comparisons, only the equation of gravitational acceleration and the presence of a DM halo should change between the Newton and MOND runs. In practice, things are more complicated.

⁴ Such velocities are unrepresentative of galaxy group and cluster environments but fully compatible with that of pairs (see e.g. Chou, Bridge & Abraham 2012).

First, the galaxy models must be set up in equilibrium. The velocities of the baryonic components must be set according to the local gravitational potential. Although MOND provides a good fit to the Newtonian rotation curve of galaxies (with an isothermal DM halo) at large distances, differences can be found in the inner regions ($\lesssim 10$ kpc), depending on the shape of the DM halo MOND replaces. Setting up stable galaxies would then require to adjust the velocity dispersions of baryons between the two cases, and thus would lead to slightly different galaxies. Such a difference could amplify over a few rotation periods and significantly alter the formation of substructures like bar(s), spiral arms and clumps. In turn, this would change the morphology of the galaxies, their intrinsic star formation activity and their response to the interaction. We circumvent this problem by using a specific galaxy model allowing for replacing the DM halo with the MOND formalism and maintaining equilibrium, as described in Section 2.5.

Secondly, the use of an adaptive grid can also introduce a bias. In RAMSES and POR, the refinement of the grid is based (among other criteria) on the number of particles in cells. Typically, a cell is refined when it contains more than 40 particles (but this number varies from simulation to simulation). Together with refinement criteria on mass and stability, this ensures that the baryonic component and the DM halo are correctly sampled on the AMR grid. However, in the MOND case, the absence of DM particles implies that the net number of particles per cell is never exactly the same as in the Newtonian run. This leads to different refinements or in other words, to different resolutions between the Newton and the MOND cases. However, the refinement is usually dictated by the total mass (stars, gas and DM) and the stability (based on the local Jeans length), both dominated by the baryons in the galactic discs. For this reason, and to avoid introducing numerical artefacts in changing the refinements criteria, we chose to keep the same strategy in the MOND and Newton cases. In practice, the differences in refinement are small, and their effects on the physical state of the galaxies are negligible, as shown in Appendix A.

2.5 Initial conditions

In their gas-free study, Candlish et al. (2015) found that the composite disc–bulge–halo model MW-A from Kuijken & Dubinski (1995) yields a very similar behaviour over several Gyr when the DM halo is replaced with QUMOND phantom DM. Therefore, such a model minimizes the differences between the MOND and Newton cases and solves the first point of Section 2.4. For this reason, we decide to use this model in all our simulations.⁵ The *N*-body components (stars and DM) are generated using the `mkkd95` tool from the `NEMO` package (Teuben 1995). Then, particles representing 10 per cent of the stellar disc mass are removed and replaced by the equivalent density distribution in gas form (in RAMSES cells). A rotation velocity is attributed to the gas cells to balance the local gravitational, turbulent and sonic pressure terms (see Chapon 2011; Renaud et al. 2013). Within a dynamical time, the gas component cools down into a thinner disc. Such an approach simplifies the initial setup in RAMSES, and is followed for both our Newton and MOND runs. In the MOND case, the DM particles are deleted. The initial conditions of our models are summarized in Table 2 (see also MW-A from Kuijken & Dubinski 1995).

⁵ The main differences with the models of Renaud et al. (2015a) are a larger disc and a more massive bulge. These changes affect the evolution and the properties of the merger, in particular the SFR, as discussed in Section 3.3.

Table 2. Initial setup of the progenitors.

<i>Gas disc (exponential)</i>	
Mass ($\times 10^9 M_\odot$)	4.2
Characteristic radius (kpc)	4.5
Truncation radius (kpc)	22.5
Characteristic height (kpc)	0.45
Truncation height (kpc)	2.25
<i>Stellar disc (exponential)</i>	
Number of particles	720 000
Mass ($\times 10^9 M_\odot$)	41.62
Characteristic radius (kpc)	4.5
Truncation radius (kpc)	22.5
Characteristic height (kpc)	0.45
Truncation height (kpc)	2.25
<i>Stellar bulge (King)</i>	
Number of particles	400 000
Mass ($\times 10^9 M_\odot$)	23.1
Radial extent (kpc)	4.5
<i>DM halo (lowered Evans)</i>	
Number of particles	600 000
Mass ($\times 10^9 M_\odot$)	261.0
Concentration	0.1
Characteristic radius (kpc)	3.6
Truncation radius (kpc)	50.0
Total mass ($\times 10^9 M_\odot$)	330.0 (Newton), 69.0 (MOND)

To verify the stability of our setups, we have run simulations of the progenitor galaxy in isolation, both in the Newton and MOND framework. Results are summarized in Appendix B.

2.6 Extended haloes

As shown by Bournaud et al. (2003), the formation of structures in tidal debris is sensitive to the size of the DM halo. Extended haloes can indeed favour the local collapse of gas and the further formation of stellar objects like clusters and TDGs. Since the MW-A model we use has initially been designed to study structures in isolated disc galaxies, the halo is truncated to a rather small value compared to typical virial radii (50 kpc instead of ≈ 200 kpc for the Milky Way; see e.g. Dehnen, McLaughlin & Sachania 2006). This provides a significant gain in computational time and memory. In studies like ours however, the tidal tails of interacting galaxies could expand to large distances, and our results could thus be affected by the artificial truncation of the halo. (In practice however, the tidal tails of our model are well within the truncation radius, see Section 3.2.)

To test the importance of the truncation, we replace the truncated lowered Evans DM halo of our models with an NFW (Navarro, Frenk & White 1997) model parametrized such that it reproduces the velocity curve in the innermost 30 kpc, and truncated at much larger radius (200 kpc). Details are presented in Appendix C. Without changing the baryonic components of our models, we run the N_A case with the extended halo. Apart from slight deviations due to the different mass of the progenitor galaxy, we find no difference in the global morphology nor in the substructures of the tidal tails.

We conclude that the truncation radius adopted here (50 kpc) is sufficiently large not to affect the formation of tidal tails nor substructures inside them. We decide to keep the truncated version (as in Table 2) to minimize the computational cost of the simulations.

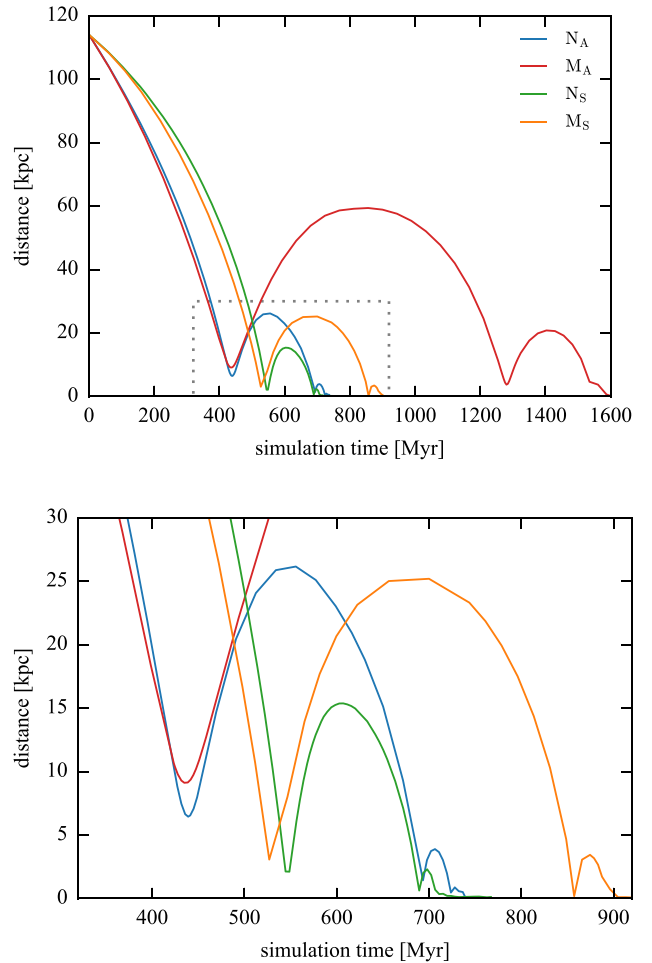


Figure 1. Evolution of the distance between the progenitor galaxies, starting at the beginning of the simulations. The bottom panel is a zoom-in. All models start with a separation of 114 kpc at $t = 0$.

3 RESULTS

3.1 Orbits

In MOND, as galaxies come close to each other, they still lie in the weak gravitational regime where MOND provides the strongest differences from the Λ CDM framework. The change in dynamical friction due to the absence of a DM halo significantly alters the exchange of orbital energy and thus the fate of the progenitors. For that reason, most of interacting galaxies in MOND have to be fly-bys, and the frequency of galaxies actually merging is much lower in MOND than in Newtonian dynamics (Nipoti et al. 2008; Kroupa 2015). For the same reason, it is already known that when the orbit allows for it, the merger of galaxies in MOND takes longer than in Λ CDM (Tiret & Combes 2008a). Furthermore, because of the absence of DM, all exchanges of orbital angular momentum are exclusively transmitted to the baryons, which also affects the disruption of the galaxies at the moment of their encounter(s) (Combes & Tiret 2010).

Fig. 1 shows the evolution of distance between the two galaxies for each simulation. We compute this as the distance between the centres of mass of the ≈ 1000 initially most gravitationally bound stellar particles of the galaxies. All simulations have been stopped after final coalescence, once the global SFR decreases back to a few $M_\odot \text{ yr}^{-1}$.

Table 3. Orbits.

Simulation	N_A	M_A	N_S	M_S
First pericentre (kpc)	6.5	9.2	2.0	3.1
Second pericentre (kpc)	1.4	3.8	0.6	0.2
Separation phase ^a (Myr)	255	845	143	331
Interaction phase ^b (Myr)	284	1149	165	368
Coalescence (normalized time ^c)	1.11	1.36	1.15	1.11

^aTime between the first and second passages.

^bTime between the first passage and the coalescence.

^cIn normalized time, the first (resp. second) pericentre passage occurs at $t' = 0$ (resp. 1). Normalized time at coalescence is thus the ratio of the durations of the interaction and the separation phases.

The differences between the Newton and MOND cases arise a few Myr before the first pericentre passage, due to the differences in dynamical friction from the DM component of the progenitors when they start to overlap. Because of the small impact parameters of the orbits used, the baryonic components also overlap adding an extra contribution to dynamical friction. In the MOND cases, only this baryonic contribution acts on the braking of the galaxies (on top of the transfer of kinetic energy and angular momentum to the tidal debris).

Because of the diversity of orbits arising from the initial velocities we choose and the differences in dynamical friction, comparing the evolution of physical quantities in our simulations can be difficult. We circumvent this by defining the time t with respect to the first pericentre passage, in the rest of the paper. Furthermore, we introduce the normalized time of interactions obtained by normalizing the time to the duration of the separation phase, i.e. the period between the first and second pericentre passages. Therefore in the following, a normalized time $t' = 0$ corresponds to the first passage, while $t' = 1$ corresponds to the second passage.

The pericentre distances and the time spans between encounters are listed in Table 3.

3.2 Morphology

Fig. 2 displays the morphologies of the models at the epochs of maximum separation between the first and second passages. In addition, Fig. 3 shows the stellar and gas density maps, and the velocity along the line of sight at the moment of the second pericentre passage.⁶

In the M_A model, the high relative velocity between the progenitors after the first passage (due to the initially high velocity and the weak effect of dynamical friction) accounts for a larger separation, and a longer period for the tidal tails to expand. At the moment of maximum separation between the progenitors, the two most distant points in the stellar tidal tails are separated by ≈ 230 kpc (≈ 160 kpc in the plane of the sky), i.e. 3.4 times more than those of N_A (2.7 times in the plane of the sky).

Whereas the progenitors are initially set up with a relative velocity twice higher in N_A than in M_S , because of the differences in dynamical friction they reach similar velocities at their respective pericentre passages (594 and 568 km s⁻¹, respectively). That is, both models have comparable relative velocities, impact parameters and masses at the time of the first encounter, which translates in the expected comparable morphology and spatial extent of the tidal tails. After the first passage, differences in dynamical friction

continue to alter the evolutions of the two models and make the separation phase of $M_S \approx 1.3$ times (76 Myr) longer than that of N_A . Yet, because of differences in angular momentum transfer, the lengths of the tidal tails remain comparable in both formalisms.

We note that the disc remnant of one of the galaxies in the M_S model (galaxy 1, located at $y > 0$ during the separation in the bottom-right panel of Fig. 2) is more concentrated than its equivalent in N_A . This galaxy is that having the strongest spin-orbit coupling (or equivalently the lowest inclination in the orbital plane, see Section 2.3), and thus the differences between the two formalisms and with the other galaxy are likely due to different responses in the angular momentum transfer during the first passage. This further affects the evolution of the system at coalescence (see below).

In M_S , the tidal tails host the formation of overdensities similar to TDGs (about 3 in each tail are visible in Fig. 3). Their associated phantom DM further favours their growth, while no such structure is visible in the stellar nor gaseous component of the N_A model. The formation of TDGs in simulations is sensitive to the resolution (Wetzstein, Naab & Burkert 2007) and the truncation of the DM halo (Bournaud et al. 2003). By conducting our comparisons at the same resolution in Newton and MOND, and by testing the formation of substructures with much more extended haloes (Appendix C), we ensure that the differences we detect have a physical origin. Observations of the Antennae galaxies report only one TDG candidate, at the tip of the southern tail (Mirabel, Dottori & Lutz 1992), but the exact nature of this structure is still questioned (Hibbard et al. 2001, see also Bournaud et al. 2004). It could be either an unbound object or a forming TDG still out of equilibrium. For the models we consider, the Newtonian framework does not allow for the formation of TDGs, while the MOND does. However we note that, in the absence of efficient shielding from the rest of the galaxy, star-forming regions in the tidal tails are more sensitive to ultraviolet radiation of extragalactic origin. A stronger radiation, e.g. at higher redshift or in a denser galactic environment, could potentially prevent the formation of TDGs and thus reduce the differences (in the young stellar component) between Newtonian and Milgromian cases in this context. Our simulations show however that the old stellar component is likely to remain more clumpy in MOND, as long as the gaseous contribution to the local gravitational potential is negligible over the stellar one. Leaving this issue aside, with the specific models considered here, the Milgromian runs tend to slightly overproduce TDGs given the absence of unambiguously defined ones in the Antennae, whilst the Newtonian runs might potentially slightly underproduce them if the observed TDG candidate turns out to be real. However, the uniqueness of our initial conditions has not been established and it is possible that other sets of parameters could reproduce the same morphology with a different number of TDGs. A much larger simulation sample including more interacting systems would be necessary to reach a clear conclusion on this particular topic.

In conclusion, despite different small-scale features ($\lesssim 1$ kpc) and without a fine tuning of the parameters, the overall morphology and kinematics of an Antennae-like system can be reproduced in the MOND framework when the progenitors have a small initial velocity to compensate for the weakness of dynamical friction. This complements the pioneer work of Tirit & Combes (2008a) who reproduced an Antennae-like morphology in MOND.⁷

⁶ The real Antennae system is thought to be observed a few Myr before the second pericentre passage, when the discs start to overlap on the eastern side, and before nuclei reach their minimum distance (Renaud et al. 2008).

⁷ Their model uses the same orbit as that of the restricted simulation of Toomre & Toomre (1972) with point-mass galaxies surrounded by massless tracer particles, thus with no dynamical friction.

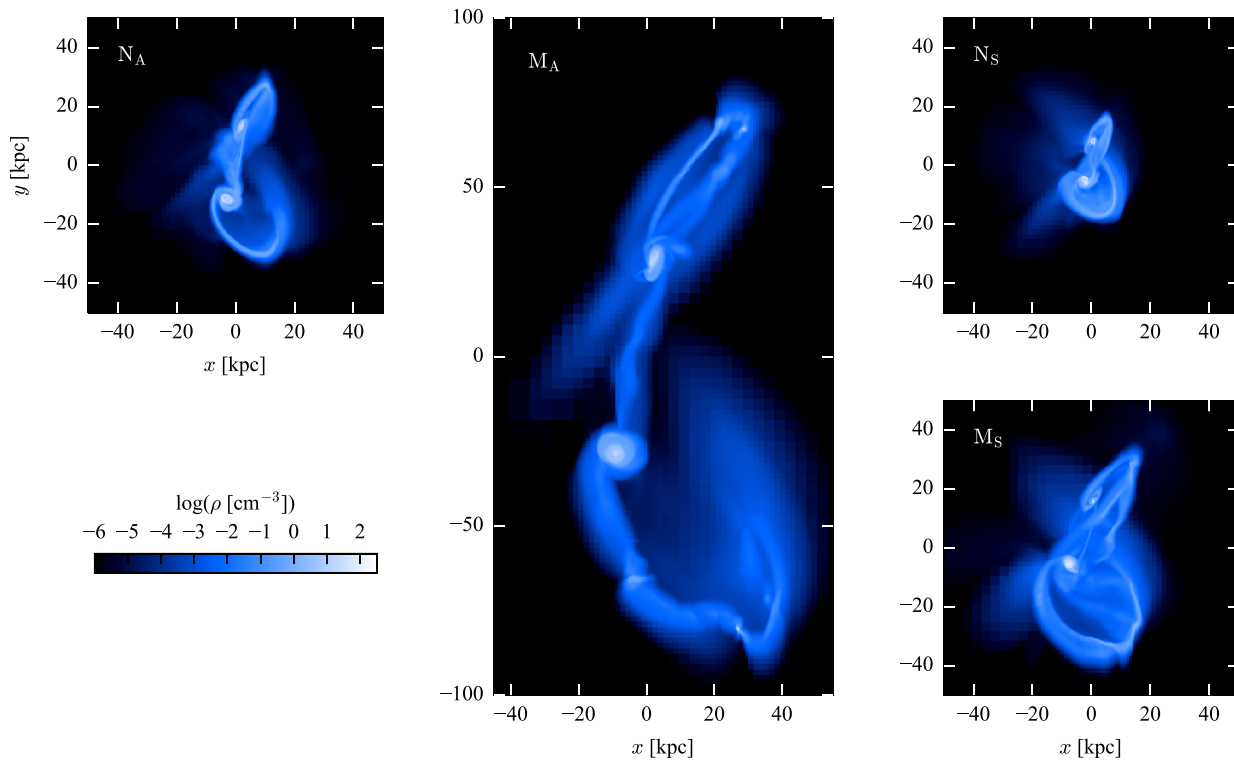


Figure 2. Map of the gas density of the models, when the galaxies reach their maximum separation after the first passage. To ease the comparison of the spatial extent of the tidal tails, all panels are at the same scale. Note the sharper and more structured features due to the baryons self-gravity in the MOND, compared to the Newtonian one.

The main differences between the Newton and MOND Antennae-like models are the concentration of the disc remnants and the presence of substructures along the tidal tails. Both these features play an important role in the star formation histories (SFHs) of the mergers, as discussed below.

3.3 Star formation

The starburst activity of interacting galaxies can be triggered by several physical processes, each having different properties and signatures.

(i) Inflow of gas towards the galactic centres (see e.g. Barnes & Hernquist 1991; Hopkins et al. 2009; Bournaud 2010). By definition, such activity is restricted to the innermost regions of the galaxies and is mostly triggered during close encounters, when one galaxy exerts negative gravitational torques on the interstellar medium (ISM) of the other, inside corotation. These torques fuel the gas inwards and thus increase its density in the innermost regions of the galaxies. This effect can also be amplified by torques generated internally by the disc structures like interaction-triggered bars and spiral arms.

(ii) Collisions between clouds and H I reservoirs (Jog & Solomon 1992; Barnes 2004). When the ISM of the progenitors overlap, it is possible that shocks make previously marginally stable gaseous structures collapse. Such an effect is limited to the overlap phase(s) and region(s) of penetrating encounters and is thus mostly active and efficient at the late stages of interactions like at coalescence. It is also possible that the overlap volume hosts cloud–cloud collisions, known to trigger intense episodes of star formation, in particular in the form of massive stars (see e.g. Loren 1976; Tan 2000; Renaud et al. 2015b).

(iii) Compression of the ISM by tides and turbulence (Renaud et al. 2014). The gravitational interaction between galaxies makes the tidal field fully compressive over kpc-scale volumes (Renaud et al. 2008, 2009). This effect is transmitted to the turbulence which develop a compressive nature, resulting in excesses of dense gas and thus an elevation of the SFR (see also Jog 2013, 2014; Mondal & Chakraborty 2015, on stability criteria). This process is active during all passages and over large volumes.

Therefore, central star formation activity is more likely to be linked to inflows. Activity in the overlap regions of the progenitors can be associated with collisions and/or compression. Finally, compression is likely responsible for triggering starburst activity over large volumes, in particular in the outskirts of galaxies and the tidal debris. In all cases and all regions, star formation can obviously still proceed in a non-enhanced way, and/or with small-scale triggers, as it does in isolated galaxies.

3.3.1 Global evolution

Figs 4 and 5 show the SFRs of all models. The differences in the ISM noticed for the isolated models (see Appendix B) induce slightly different intrinsic SFRs. For our Newtonian models, the total SFR in the isolation phase (i.e. before any sign of interaction) is $1.4 M_{\odot} \text{ yr}^{-1}$, while it is $0.9 M_{\odot} \text{ yr}^{-1}$ in the MOND cases. We also note that the MOND models have a slightly more spatially extended star formation in isolation: 3 per cent of their stars form beyond 5 kpc, while this fraction is less than 1 per cent in the Newtonian cases.

All simulations yield an increase of their star formation activity at the times of all their pericentre passages, but with remarkably different amplitudes. Although the relative importance of the

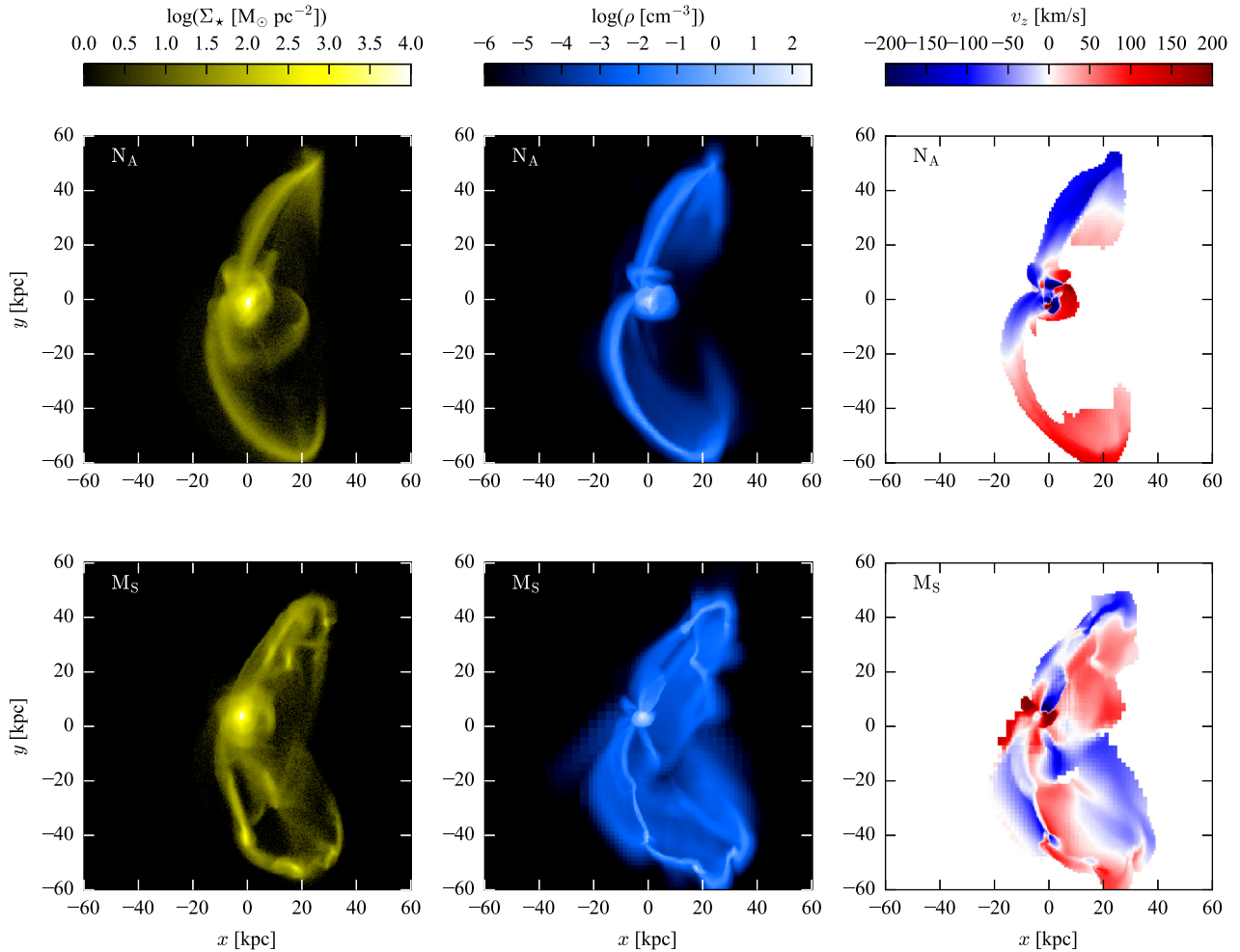


Figure 3. Map of the stellar surface density (left), gas density (centre) and velocity along the line of sight (right) of the N_A (top) and M_S (bottom), at the second pericentre passage. For the sake of clarity, only the velocity field in regions of high gas density is shown.

physical triggers responsible for these bursts varies from one model to the next, any enhancement of the star formation activity is stronger when its trigger acts for a long period. Therefore, systems with low relative velocities (N_S and, to a lower extent, M_S) harbour a stronger burst of star formation. The fastest model (i.e. with the highest initial velocity and the least dynamical friction) M_A yields the lowest burst at the first passage. In short, the intensity of the first burst in our models decreases with increasing velocity of the progenitors. This result seems to contradict the trend proposed by Di Matteo et al. (2007, their fig. 22), but we note that their sample of direct fly-bys for late-type galaxies (which corresponds to all our cases at first passage) does not yield a clear trend between the intensity of starburst and the relative velocity of the progenitors in the range covered by our models ($\approx 500\text{--}600\text{ km s}^{-1}$ at pericentre).

In all models, triggered star formation starts a few Myr before the pericentre passage itself, and over larger volumes than during the isolation phase. This is likely the signature of the long-range effect of tidal compression. Just after the pericentre, the fraction of SFR in the outer regions drops, indicating that gas has been fuelled towards the nuclei and forms stars. The delay between the trigger (torques) and the resulting activity (SFR enhancement) is slightly shorter in the Newton runs than in our Milgromian models.

After the pericentre passage, as the galaxies start to separate, the tidal debris is ejected. The presence of a massive bulge in our models helps maintaining a relatively high SFR during the separation phase, while it rapidly decreases down to almost its pre-interaction value in the simulation of Renaud et al. (2015a) who used the same orbit as N_A but different progenitor galaxies. Interestingly, the maximum SFR before coalescence corresponds to the pericentre passage for the Newton models, while it is reached later during the separation in the Milgromian runs, and is associated with formation outside the disc remnants, i.e. in the weak gravitational regime, exactly where differences between Milgromian and Newtonian dynamics are expected (see Section 3.3.2).

At the time of the second passage and coalescence, the configuration of the first passage repeats itself. Because of the orbital energy transferred to the tidal debris, the impact parameter and relative velocities are however smaller than before (see Fig. 1) and the above-mentioned effects (torques, collisions in overlaps) get even stronger than at the first encounter. Thus, most of the star formation takes place in the central regions. In the M_A model, the second passage is very similar to the first one, both having relative velocities of $\approx 500\text{ km s}^{-1}$, leading to a weak enhancement of the SFR (to about $5\text{ M}_\odot\text{ yr}^{-1}$).

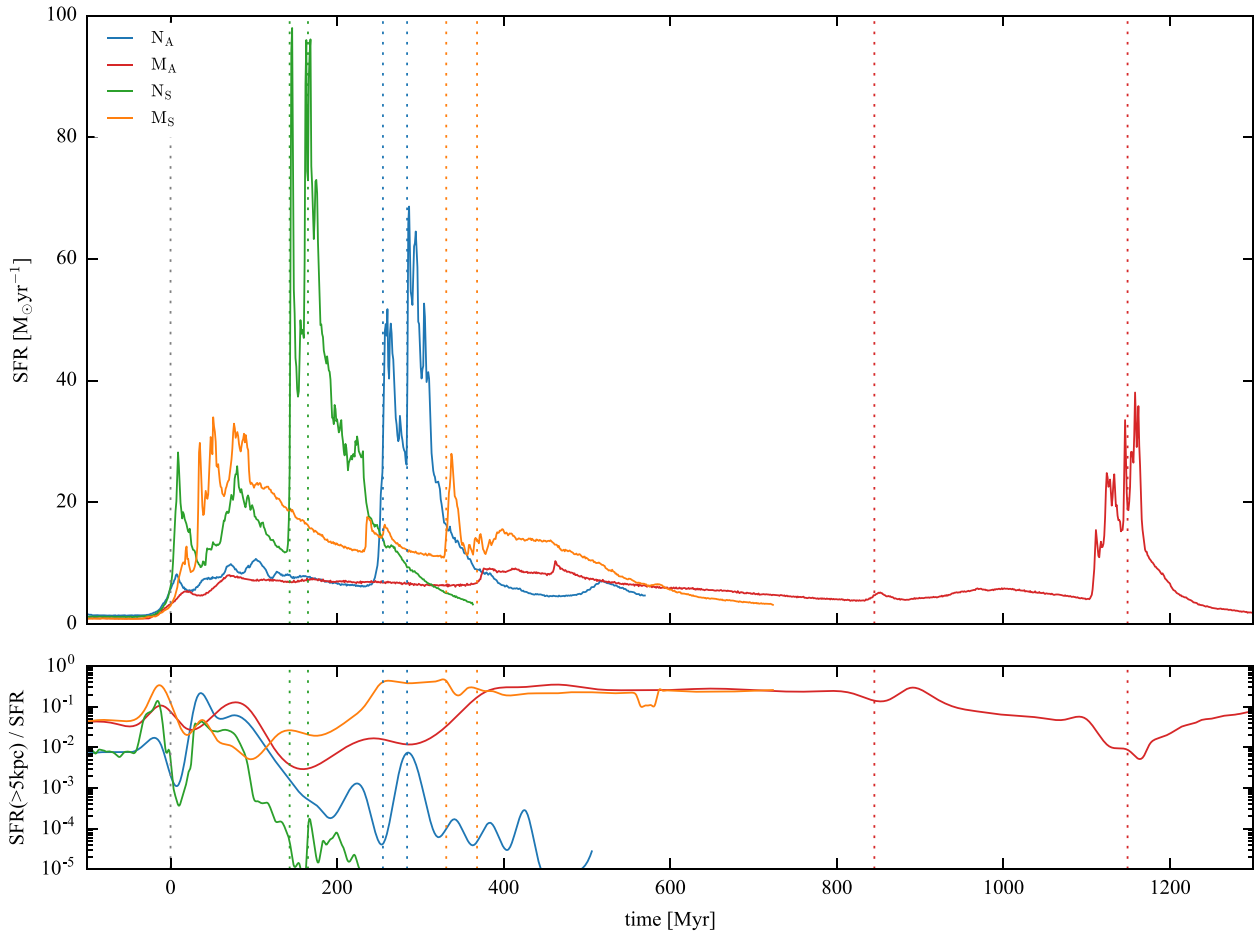


Figure 4. SFR (top) and fraction of the SFR further than 5 kpc from the galaxy centres (bottom). A large fraction denotes star formation activity in the tidal debris. $t = 0$ corresponds to the first pericentre passage. Vertical lines indicate the first and second pericentre passages, and the final coalescence.

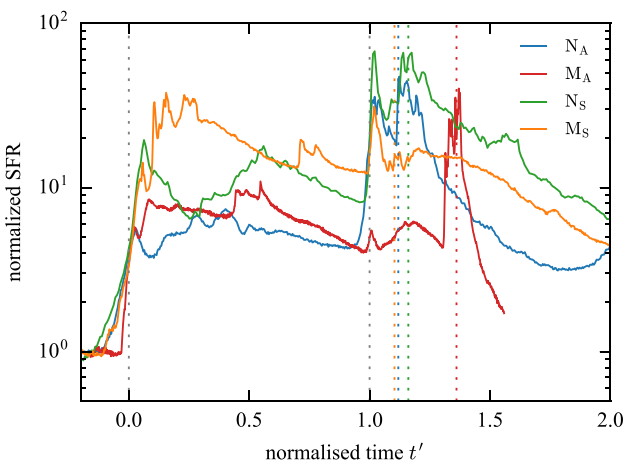


Figure 5. SFR normalized to its value before the interaction, as a function of time normalized to the separation period (i.e. 0 on the x -axis corresponds to first pericentre passage, while 1 indicates the second passage). The vertical lines indicate final coalescence, corresponding to the third encounter for these models.

At coalescence, all models host a burst of star formation activity, which lasts for about ≈ 100 Myr. The intensity of this burst depends, as before, on the orbital configuration, but also on the amount of gas left in the reservoir after the precedent bursts.

3.3.2 Star formation in the tidal debris

For all models, enhanced star formation activity starts a few Myr before the first pericentre passage, in the outer regions. This corresponds to the distant effect of compressive tides acting on large scales over the progenitors. The tidal debris, namely a bridge connecting the two progenitors and tails on the other sides, forms and expands during the separation phase. In the Newtonian cases, the gas clouds in the tidal debris are destroyed by dynamical processes like extensive (classical, destructive) tides and shear (Renaud et al. 2015a). Star formation in the outer regions is thus quenched.

The ISM of the tidal bridges connecting the two galaxies is more fragmented than the tails (Fig. 2), and hosts the formation of a few star clusters (see also Renaud et al. 2015a). There, the inner regions of the two DM haloes overlap, as opposed to the tails where the more distant halo of the companion has a negligible effect.

In the MOND cases however, the tidal debris generates its own ‘phantom’ potential well which favours the gathering of gaseous overdensities, leading to star formation in TDG-like objects (Figs 2 and 3). These formation events are visible in the total SFR (Fig. 4) as peaks during the separation phase (e.g. at $t \approx 240$ Myr in M_S , and $t \approx 460$ Myr in M_A). At these moments, the fraction of SFR in the outer regions increases significantly (although the activity in the discs also continues). This constitutes a major difference between the Newtonian and Milgromian families of models.

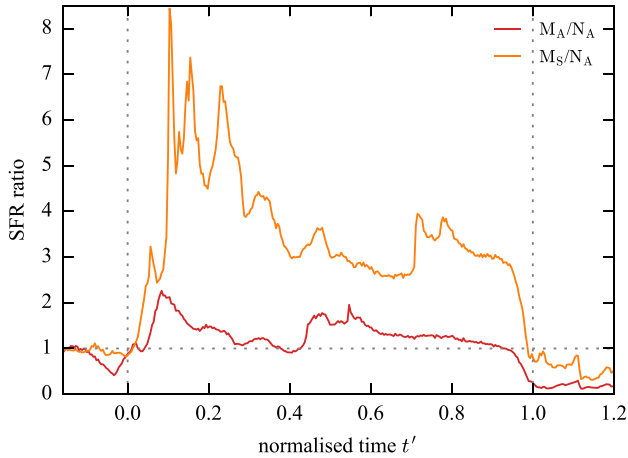


Figure 6. Ratio of the SFRs between the M_A and N_A , and the M_S and N_A simulations, i.e. our two sets of most direct comparisons. As in Fig. 5, time is normalized to the separation period. Note that, in this time unit, N_A , M_A and M_S models reach coalescence at 1.12, 1.37 and 1.13, respectively.

3.3.3 Comparisons between Newtonian and Milgromian gravitations

In this section, we focus on the differences between our most comparable models, namely N_A and M_A which share the same initial conditions, and N_A and M_S which yield comparable morphologies.

Fig. 6 shows the ratios of SFR between our models. The increase of the SFR at the pericentre passages is fairly independent of the gravitational paradigm (N_A and M_A) but varies with the velocity of the encounter (N_A and M_S). We note from Fig. 5 that the first stages in the star formation burst a few Myr before the pericentre itself ($-0.1 \lesssim t' \lesssim 0$) are remarkably similar in the N_A and M_S runs. Since this phase is mostly associated with tidal compression, we conclude that the Milgromian gravitation produces here a similar effect as the DM halo at these epochs. The same situation is found again at the moment of the second encounter passages ($t' \approx 1$), despite very different evolutions in between.

As noted before, galaxies in the Milgromian framework are more efficient at forming stars in between the pericentre passages than the Newtonian case. A significant fraction of this activity (up to 40 per cent) is located in the tidal debris, but the majority happens in the disc remnants. At this stage, star formation takes place in the fragmented ISM and because of tidal compression. However at large separations, the Newtonian galaxies lie in the outer regions of the DM haloes of each other, where tidal compression (if/where/when it exists) is weaker than in the inner halo (Renaud et al. 2009). The Milgromian formalism is then more efficient at triggering star formation in the discs, in a comparable fashion as the situation in tidal debris (see Section 3.3.2).

The Newtonian case is then forming more stars after the second passage ($t' > 1$) than both Milgromian models. This activity takes place in the disc remnants of the galaxies, in the central ≈ 5 kpc, as visible in Fig. 4. Gas inflows driven by gravitational torques favour such star formation. The differences between the Newtonian and Milgromian models come from the amount of gas available in the volume affected by these torques, which depends on the angular momentum transfer during the first passage (see the different disc morphologies and densities in Fig. 2). Even though torques of comparable amplitude exist in all models, the Newtonian one has retained more gas to fuel the nuclear star formation activity than the Milgromian cases. A similar effect can be noted in

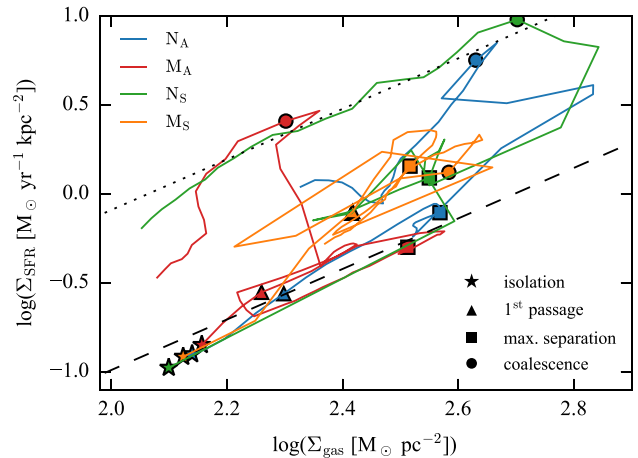


Figure 7. Evolution of the gas and SFR surface densities in the central 2 kpc of the galaxy(ies). Before coalescence, the points are the average of the two progenitors. The dashed and dotted lines indicate the observed sequences of discs and of starbursts, respectively, as in Daddi et al. (2010a, see also Genzel et al. 2010). Stars, triangles, squares and dots mark the moments before the interaction, during the first starburst (i.e. a few Myr after the first passage), at the maximum separation of the galaxies, and during the second starburst (about at the coalescence).

retrograde encounters where momentum transfer is very inefficient at the first passage, which maintains a large gas reservoir for the later encounters (see e.g. Duc & Renaud 2013).

Because of the longer interaction period of the Milgromian runs, over the entire interaction (i.e. from 40 Myr before the first passage until the end of the simulations, when the SFRs have approximately reached back their pre-interaction level), the total mass of stars formed is $6.7 \times 10^9 M_\odot$ in the N_A run, and 30, 13 and 37 per cent more in the M_A , N_S and M_S models, respectively. However, the spread in formation epochs makes such differences difficult to detect observationally, as discussed in the next section.

To conclude, significant differences exist in the star formation activities of the Newtonian and Milgromian models, in particular in between the pericentre passages. Part of these differences corresponds to star formation in the tidal debris, as noted in Fig. 4. Such activity does not exist in the Newtonian model, which is less prone to form TDGs in this simulated encounter as mentioned earlier.

3.4 Observational diagnostics

Fig. 7 shows the evolution of the models in the Schmidt–Kennicutt diagram, linking the surface density of gas Σ_{gas} to the surface density of SFR Σ_{SFR} . By construction, all models start close to the regime of discs noted by Daddi et al. (2010a) and Genzel et al. (2010). At the first passage, they become more compact and more efficient at forming stars, as discussed above. The maximum efficiencies are reached around the coalescence. The Newtonian models reach higher densities than the MOND ones, again because of differences in angular momentum transfer away from the central regions.

Following Fensch et al. (in preparation), we define the starburstiness parameter as the ratio of the measured surface density of SFR (Σ_{SFR}) to the value it would have if the galaxies were on the disc sequence for the measured surface density of gas (Σ_{gas}). We consider the galaxies as starbursting when their starburstiness exceeds 4 (see e.g. Schreiber et al. 2015). Fig. 8 shows the evolution of the starburstiness for the four models. Galaxies are considered as

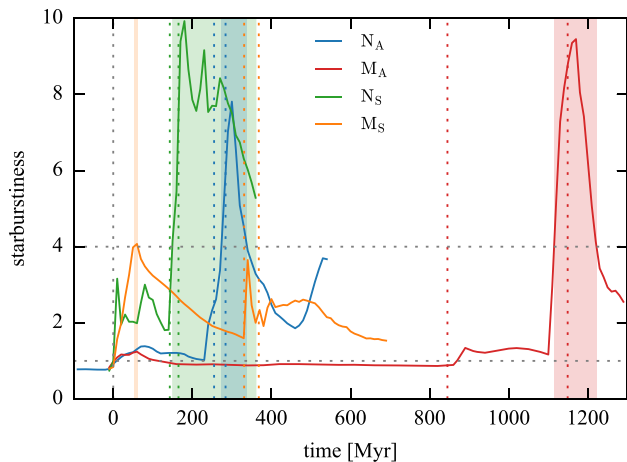


Figure 8. Evolution of the starburstiness parameter, defined as the ratio between the measured Σ_{SFR} and that of the disc sequence for the corresponding measured Σ_{gas} . $t = 0$ corresponds to the first pericentre passage. The shaded areas mark the starbursting phases of the interactions (i.e. with a starburstiness above 4).

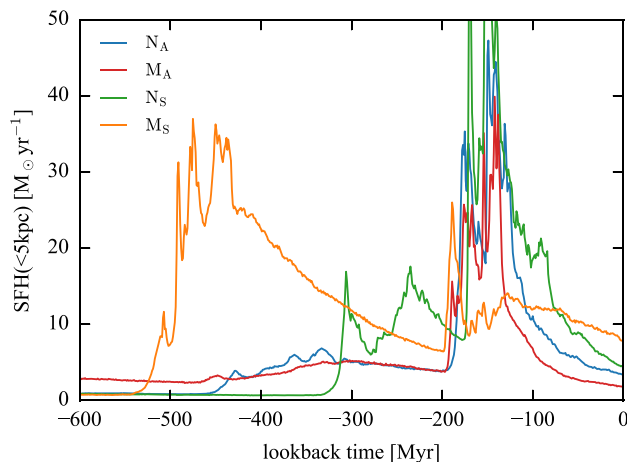


Figure 9. SFH of the stars in the central 5 kpc. To adopt an observational perspective, here the time axis is reversed, and $t = 0$ corresponds to 150 Myr after the beginning of final coalescence.

starbursting for 65, 108 and 9 Myr for the N_A , M_A and M_S runs, respectively, and for more than 213 Myr in the N_S run where we stopped the simulation before the system leaves this regime. Since the starburst regime is only reached (if it is at all) at the pericentres and not during the entire interaction, the longer interaction periods of the MOND cases do not imply a longer duration of the starburst phase. Therefore, for comparable morphologies, the Newtonian systems are more likely to be observed in a starburst phase like (ultra-)luminous infrared galaxies (LIRGs, ULIRGs; e.g. Kennicutt 1998) than their Milgromian equivalent.

Finally, from an observational perspective, Fig. 9 shows the SFH of the merger remnants, i.e. the histogram of ages for the stars detected in the central 5 kpc, 150 Myr after the beginning of final coalescence. As in observations, most of the stars form in this volume, but some have been accreted from the tidal debris. The recent SFH (i.e. during coalescence, less than 200 Myr in the past) is comparable for all models with differences in the amplitude of the burst as discussed above. However, the differences in orbital period and in the burst associated with the first encounter provide a strong

dichotomy between the Newtonian and Milgromian models. When comparing models with similar morphologies (N_A and M_S), we note that the SFH in MOND is more extended and yields a stronger burst in the ≈ 300 – 600 Myr range.

To conclude, interacting galaxies in MOND host extended star formation in time and space, which makes them less likely to be detected as starbursts as their Newton equivalents.

4 SUMMARY AND DISCUSSION

Using hydrodynamical simulations, we present a comparison as directly as possible between interacting galaxies within the Newtonian and MOND frameworks. Our main results are as follows.

(i) Replacing DM with a MOND formalism induces differences in dynamical friction and angular momentum transfer during the galactic encounters. As a result, the tidal debris spans a much larger volume in Milgromian models than in Newtonian ones for a given set of initial conditions.

(ii) Comparable merger morphologies can still be obtained in Newtonian and Milgromian frameworks. We find that merely tuning down the relative velocity of the progenitors to balance the weak dynamical friction in MOND and reach comparable velocities at the first pericentre in both frameworks provides a reasonable correspondence between the models, at least on large scales. Thus, without a fine tuning of the parameters, the overall morphology and kinematics of an Antennae-like system can be reproduced in the MOND framework. This complements the pioneer work of Tirit & Combes (2008a) who reproduced an Antennae-like morphology in MOND for the first time.

(iii) While the global morphology and kinematic structure can be reproduced with MOND, small-scale differences from the Newtonian model exist: mainly, the formation of stellar and gaseous clumps along the tidal tails and the spatial extent of the discs at the moment of the second encounter. By generating their own phantom DM and thus a deep potential well on small scales ($\lesssim 1$ kpc), the tidal tails favour the collapse of dense gas structures and thus the formation of stars. A significant fraction of the total star formation occurs in the tidal debris in Milgromian gravitation, while such activity is negligible in the Newtonian case. The Milgromian models thus lead to significantly more substructures in the tidal tails than the Newtonian models.

(iv) The resulting star formation activity is thus significantly more extended in space and time in Milgromian than in Newtonian gravity.

The results presented here originate from a handful of orbital configurations and for only one galaxy model. Generalizing our conclusions to other systems, over the broad range of observed parameters (mass, mass ratio, relative velocity, inclination, spin-orbit coupling, etc.), would require a more complete survey of simulations. Such a study is proposed in the Newtonian framework by the GALMER project (Di Matteo et al. 2007), where hundreds of configurations are explored. An equivalent in Milgromian gravitation would allow us to broaden our conclusions on the response of Milgromian galaxies to interactions and mergers.

In particular, we expect that the SFH of retrograde encounters (i.e. orbital and disc angular momenta being anti-aligned) in MOND would be more similar to their Newtonian counterparts than for the cases we presented here. During the first pericentre passage of a retrograde encounter, the orbital angular momentum is inefficiently transferred to the discs and tidal features are less pronounced and shorter than in prograde cases (see e.g. Duc & Renaud 2013). As a

result, a large gas reservoir remains available at the late stages of the interaction, in particular at coalescence, where cloud–cloud collisions and nuclear inflows can efficiently trigger an intense episode of star formation. In that case, the SFH would only yield one major peak at coalescence. The main differences between Newtonian and Milgromian models noted above in terms of space and time extent of star formation during the separation phase might not be found in a retrograde encounter.

Our results show that the observational detection or the absence of a sustained star formation activity in interacting galaxies, in particular over large volumes in tidal debris, would provide a strong hint on the nature of gravitation on galactic scales, as a complement to the study of rotation curves for isolated discs. Even when focusing on the merger remnant only, we show that the Milgromian paradigm favours a long and approximately continuous episode of star formation, starting at the first encounter and ending a few 100 Myr after coalescence. Oppositely, the Newtonian formalism rather supports distinct bursts of star formation at high efficiency, associated with the close encounters (see also Di Matteo et al. 2007). Providing that observational techniques would be able to tell apart these two SFHs with sufficient confidence in the required age range (≈ 200 Myr–1 Gyr; see Lançon 2001; Maraston et al. 2001, 2004; Simones et al. 2014, on the difficulties of dating post-starburst episodes), such interacting systems could be used as clear diagnostics to test different paradigms for gravitation on galactic scales.

ACKNOWLEDGEMENTS

We thank the referee for constructive comments which helped to improve the manuscript. This work was granted access to the PRACE Research Infrastructure resource *Curie* hosted at the TGCC (France). FR acknowledges support from the European Research Council through grant ERC-StG-335936. BF acknowledges the financial support from the *Programme Investissements d'Avenir* (PIA) of the *Initiative d'Excellence* (IdEx) programme from Université de Strasbourg.

REFERENCES

- Angus G. W., van der Heyden K. J., Famaey B., Gentile G., McGaugh S. S., de Blok W. J. G., 2012, *MNRAS*, 421, 2598
- Armus L., Heckman T., Miley G., 1987, *AJ*, 94, 831
- Arp H., 1966, *Atlas of Peculiar Galaxies*. California Inst. Technology, Pasadena
- Barnes J. E., 2004, *MNRAS*, 350, 798
- Barnes J. E., Hernquist L. E., 1991, *ApJ*, 370, L65
- Barnes J. E., Hernquist L., 1992, *ARA&A*, 30, 705
- Bate M. R., Bonnell I. A., 2005, *MNRAS*, 356, 1201
- Bonnell I. A., Dobbs C. L., Smith R. J., 2013, *MNRAS*, 430, 1790
- Boquien M., Lisenfeld U., Duc P.-A., Braine J., Bournaud F., Brinks E., Charmandaris V., 2011, *A&A*, 533, A19
- Bournaud F., 2010, in Smith B., Higdon J., Higdon S., Bastian N., eds, *ASP Conf. Ser. Vol. 423, Galaxy Wars: Stellar Populations and Star Formation in Interacting Galaxies*. Astron. Soc. Pac., San Francisco, p. 177
- Bournaud F., Duc P.-A., Masset F., 2003, *A&A*, 411, L469
- Bournaud F., Duc P.-A., Amram P., Combes F., Gach J.-L., 2004, *A&A*, 425, 813
- Boylan-Kolchin M., Bullock J. S., Kaplinghat M., 2011, *MNRAS*, 415, L40
- Brada R., Milgrom M., 1999, *ApJ*, 519, 590
- Candlish G. N., Smith R., Fellhauer M., 2015, *MNRAS*, 446, 1060
- Chapon D., 2011, PhD thesis, Université Paris 7
- Chien L.-H., Barnes J. E., 2010, *MNRAS*, 407, 43
- Chien L.-H., Barnes J. E., Kewley L. J., Chambers K. C., 2007, *ApJ*, 660, L105
- Chou R. C. Y., Bridge C. R., Abraham R. G., 2012, *ApJ*, 760, 113
- Combes F., Tiret O., 2010, in Alimi J.-M., Fuözfa A., eds, *AIP Conf. Ser. Vol. 1241, Invisible Universe*. Am. Inst. Phys., New York, p. 154
- Courty S., Alimi J. M., 2004, *A&A*, 416, 875
- Daddi E. et al., 2010a, *ApJ*, 713, 686
- Daddi E. et al., 2010b, *ApJ*, 714, L118
- de Blok W. J. G., 2010, *Adv. Astron.*, 2010, 789293
- Dehnen W., McLaughlin D. E., Sachania J., 2006, *MNRAS*, 369, 1688
- Di Matteo P., Combes F., Melchior A., Semelin B., 2007, *A&A*, 468, 61
- Dubinski J., Mihos J. C., Hernquist L., 1996, *ApJ*, 462, 576
- Dubois Y., Teyssier R., 2008, *A&A*, 477, 79
- Duc P.-A., Renaud F., 2013, in Souchay J., Mathis S., Tokieda T., eds, *Lecture Notes in Physics, Vol. 861, Tides in Astronomy and Astrophysics*. Springer Verlag, Berlin, p. 327
- Ellison S. L., Mendel J. T., Patton D. R., Scudder J. M., 2013, *MNRAS*, 435, 3627
- Famaey B., Binney J., 2005, *MNRAS*, 363, 603
- Famaey B., McGaugh S. S., 2012, *Living Rev. Relativ.*, 15, 10
- Flores H., Hammer F., Fouquet S., Puech M., Kroupa P., Yang Y., Pawlowski M., 2016, *MNRAS*, 457, L14
- Gentile G., Famaey B., Combes F., Kroupa P., Zhao H. S., Tiret O., 2007, *A&A*, 472, L25
- Genzel R. et al., 2010, *MNRAS*, 407, 2091
- Hees A., Famaey B., Angus G. W., Gentile G., 2016, *MNRAS*, 455, 449
- Hennebelle P., Chabrier G., 2008, *ApJ*, 684, 395
- Hibbard J. E., Mihos J. C., 1995, *AJ*, 110, 140
- Hibbard J. E., van der Hulst J. M., Barnes J. E., Rich R. M., 2001, *AJ*, 122, 2969
- Hopkins P. F., 2013, *MNRAS*, 430, 1880
- Hopkins P. F., Cox T. J., Younger J. D., Hernquist L., 2009, *ApJ*, 691, 1168
- Houk J. R., Schneider D. P., Danielson G. E., Neugebauer G., Soifer B. T., Beichman C. A., Lonsdale C. J., 1985, *ApJ*, 290, L5
- Ibata R. A. et al., 2013, *Nature*, 493, 62
- Ibata N. G., Ibata R. A., Famaey B., Lewis G. F., 2014, *Nature*, 511, 563
- Jog C. J., 2013, *MNRAS*, 434, L56
- Jog C. J., 2014, *AJ*, 147, 132
- Jog C. J., Solomon P. M., 1992, *ApJ*, 387, 152
- Kennicutt R. C., 1998, *ApJ*, 498, 541
- Kroupa P., 2015, *Can. J. Phys.*, 93, 169
- Kroupa P., Theis C., Boily C. M., 2005, *A&A*, 431, 517
- Kroupa P. et al., 2010, *A&A*, 523, A32
- Kuijken K., Dubinski J., 1995, *MNRAS*, 277, 1341
- Lançon A., 2001, in Tacconi L., Lutz D., eds, *Starburst Galaxies: Near and Far*. Springer-Verlag, Heidelberg, p. 231
- Lelli F., Fraternali F., Verheijen M., 2013, *MNRAS*, 433, L30
- Lelli F. et al., 2015, *A&A*, 584, A113
- Lelli F., McGaugh S. S., Schombert J. M., 2016a, *ApJ*, 816, L14
- Lelli F., McGaugh S. S., Schombert J. M., Pawlowski M. S., 2016b, *ApJ*, 827, L19
- Llinares C., Knebe A., Zhao H., 2008, *MNRAS*, 391, 1778
- Loren R. B., 1976, *ApJ*, 209, 466
- Lüghausen F., Famaey B., Kroupa P., 2015, *Can. J. Phys.*, 93, 232
- McGaugh S. S., Schombert J. M., Bothun G. D., de Blok W. J. G., 2000, *ApJ*, 533, L99
- Mac Low M.-M., Klessen R. S., 2004, *Rev. Mod. Phys.*, 76, 125
- Maraston C., Kissler-Patig M., Brodie J. P., Barmby P., Huchra J. P., 2001, *A&A*, 370, 176
- Maraston C., Bastian N., Saglia R. P., Kissler-Patig M., Schweizer F., Goudfrooij P., 2004, *A&A*, 416, 467
- Metz M., Kroupa P., Jerjen H., 2007, *MNRAS*, 374, 1125
- Metz M., Kroupa P., Libeskind N. I., 2008, *ApJ*, 680, 287
- Mihos J. C., Dubinski J., Hernquist L., 1998, *ApJ*, 494, 183
- Milgrom M., 1983, *ApJ*, 270, 365
- Milgrom M., 2010, *MNRAS*, 403, 886
- Milgrom M., 2016, preprint ([arXiv:1607.05103](https://arxiv.org/abs/1607.05103))

- Mirabel I. F., Dottori H., Lutz D., 1992, *A&A*, 256, L19
Mondal S., Chakraborty S., 2015, *MNRAS*, 450, 1874
Moreno J., Bluck A. F. L., Ellison S. L., Patton D. R., Torrey P., Moster B. P., 2013, *MNRAS*, 436, 1765
Navarro J. F., Frenk C. S., White S. D. M., 1997, *ApJ*, 490, 493
Nipoti C., Londrillo P., Ciotti L., 2007, *ApJ*, 660, 256
Nipoti C., Ciotti L., Binney J., Londrillo P., 2008, *MNRAS*, 386, 2194
Oman K. A. et al., 2015, *MNRAS*, 452, 3650
Papastergis E., Giovanelli R., Haynes M. P., Shankar F., 2015, *A&A*, 574, A113
Papastergis E., Adams E. A. K., van der Hulst J. M., 2016, *A&A*, 593, A39
Pawlowski M. S., Pflamm-Altenburg J., Kroupa P., 2012, *MNRAS*, 423, 1109
Pawlowski M. S., Famaey B., Merritt D., Kroupa P., 2015, *ApJ*, 815, 19
Privon G. C., Barnes J. E., Evans A. S., Hibbard J. E., Yun M. S., Mazzarella J. M., Armus L., Surace J., 2013, *ApJ*, 771, 120
Quinn P. J., Hernquist L., Fullagar D. P., 1993, *ApJ*, 403, 74
Read J. I., Iorio G., Agertz O., Fraternali F., 2016, *MNRAS*, 462, 3628
Renaud F., Boily C. M., Fleck J.-J., Naab T., Theis C., 2008, *MNRAS*, 391, L98
Renaud F., Boily C. M., Naab T., Theis C., 2009, *ApJ*, 706, 67
Renaud F. et al., 2013, *MNRAS*, 436, 1836
Renaud F., Bournaud F., Kraljic K., Duc P.-A., 2014, *MNRAS*, 442, L33
Renaud F., Bournaud F., Duc P.-A., 2015a, *MNRAS*, 446, 2038
Renaud F. et al., 2015b, *MNRAS*, 454, 3299
Saintonge A. et al., 2012, *ApJ*, 758, 73
Sanders D. B., Mirabel I. F., 1996, *ARA&A*, 34, 749
Schombert J. M., Wallin J. F., Struck-Marcell C., 1990, *AJ*, 99, 497
Schreiber C. et al., 2015, *A&A*, 575, A74
Scudder J. M., Ellison S. L., Torrey P., Patton D. R., Mendel J. T., 2012, *MNRAS*, 426, 549
Simones J. E. et al., 2014, *ApJ*, 788, 12
Smith B. J. et al., 2010, in Smith B., Higdon J., Higdon S., Bastian N., eds, *ASP Conf. Ser. Vol. 423, Galaxy Wars: Stellar Populations and Star Formation in Interacting Galaxies*. Astron. Soc. Pac., San Francisco, p. 227
Springel V., White S. D. M., 1999, *MNRAS*, 307, 162
Struck C., 1999, *Phys. Rep.*, 321, 1
Tan J. C., 2000, *ApJ*, 536, 173
Teuben P., 1995, in Shaw R. A., Payne H. E., Hayes J. J. E., eds, *ASP Conf. Ser. Vol. 77, Astronomical Data Analysis Software and Systems IV*. Astron. Soc. Pac., San Francisco, p. 398
Teyssier R., 2002, *A&A*, 385, 337
Teyssier R., Chapon D., Bournaud F., 2010, *ApJ*, 720, L149
Teyssier R., Pontzen A., Dubois Y., Read J. I., 2013, *MNRAS*, 429, 3068
Tiret O., Combes F., 2007, *A&A*, 464, 517
Tiret O., Combes F., 2008a, in Funes J. G., Corsini E. M., eds, *ASP Conf. Ser. Vol. 396, Formation and Evolution of Galaxy Disks*. Astron. Soc. Pac., San Francisco, p. 259
Tiret O., Combes F., 2008b, *A&A*, 483, 719
Toomre A., Toomre J., 1972, *ApJ*, 178, 623
Wetzstein M., Naab T., Burkert A., 2007, *MNRAS*, 375, 805

APPENDIX A: ADAPTIVE REFINEMENT

As mentioned in Section 2.4, since *one* of the refinement criteria used by *RAMSES* and *POR* is based on the number of particles per cell, the absence of DM particles in the MOND simulations introduces differences in the effective resolution of our Newtonian and Milgromian runs.

To quantify this effect as accurately as possible, we start a simulation of an isolated galaxy (as described in Section 2.5) in the Newtonian framework, and evolve it for 100 Myr. At this point, we duplicate the simulation and keep one copy running. In the other one, we remove the DM particles and switch on the MOND Poisson solver. Fig. A1 shows the occupation of the refinement levels

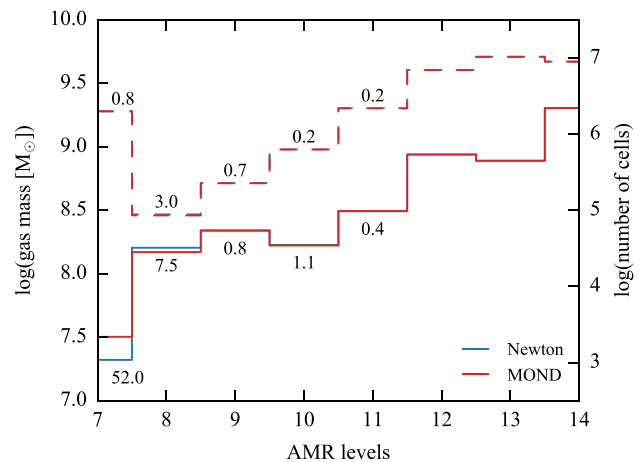


Figure A1. Gas mass (solid lines, left axis) and number of cells (dashed lines, right axis) per refinement levels in simulations of an isolated galaxy (see the text for details). The numbers indicate the percentage relative difference, i.e. $100 \times (1 - \text{MOND}/\text{Newton})$. For the finest levels, the relative difference (not shown) is smaller than 0.1 per cent. Since the simulation volume for these tests spans $(100 \text{ kpc})^3$, the size of the AMR cells is $100/2^{\text{level}}$ kpc, i.e. from 781 pc at level 7 to 6 pc at level 14. For all the simulations presented here, the minimum refinement level is 7, meaning that the entire simulation volume is at least resolved at this level.

(only for the gas denser than 10^{-3} cm^{-3}) in both runs, after another 50 Myr of evolution.

The differences mainly arise in the coarsest levels (7 and 8), i.e. the diffuse interstellar and intergalactic media. This corresponds to volumes in the outermost regions of the galaxy, i.e. where the presence or absence of DM is the most relevant and dominates the refinement strategy. For finer levels, the gaseous and stellar components dominate the local density and thus the refinement is mostly governed by criteria mildly dependent on the presence of DM (the local Jeans length and the gas mass). Therefore, the differences between the two runs remain small.

APPENDIX B: ISOLATED GALAXIES

To verify the intrinsic stability of our galaxy models, we run them in isolation at the resolution of 50 pc and without star formation. Figs B1 and B2 show the evolution of the face-on and edge-on morphology for the stellar and gaseous components. Fig. B3 displays the evolution of the velocity curves, computed using the net gravitational potential.

At the beginning (recall Section 2.5), the gas component cools down into a thin disc. Once this stage has been reached, the ISM remains approximately the same. Because of this initial phase, we start all galaxy–galaxy models at least 100 Myr before the first passage, allowing them to reach a relaxed stage before the interaction.

We note that the stellar component remains fairly stable over the duration of the integration (600 Myr, i.e. several rotation periods), except for a mild enhancement of the density in the central ≈ 2 kpc, for both models, and the formation of a spiral instability in the Milgromian case. The gas follows the stellar potential well and also forms spiral arms. Although this introduces differences between the two models, their azimuthal averaged density profiles remain comparable.

Fig. B4 shows the relative difference between the circular velocities of the isolated galaxy from the phantom DM in the MOND

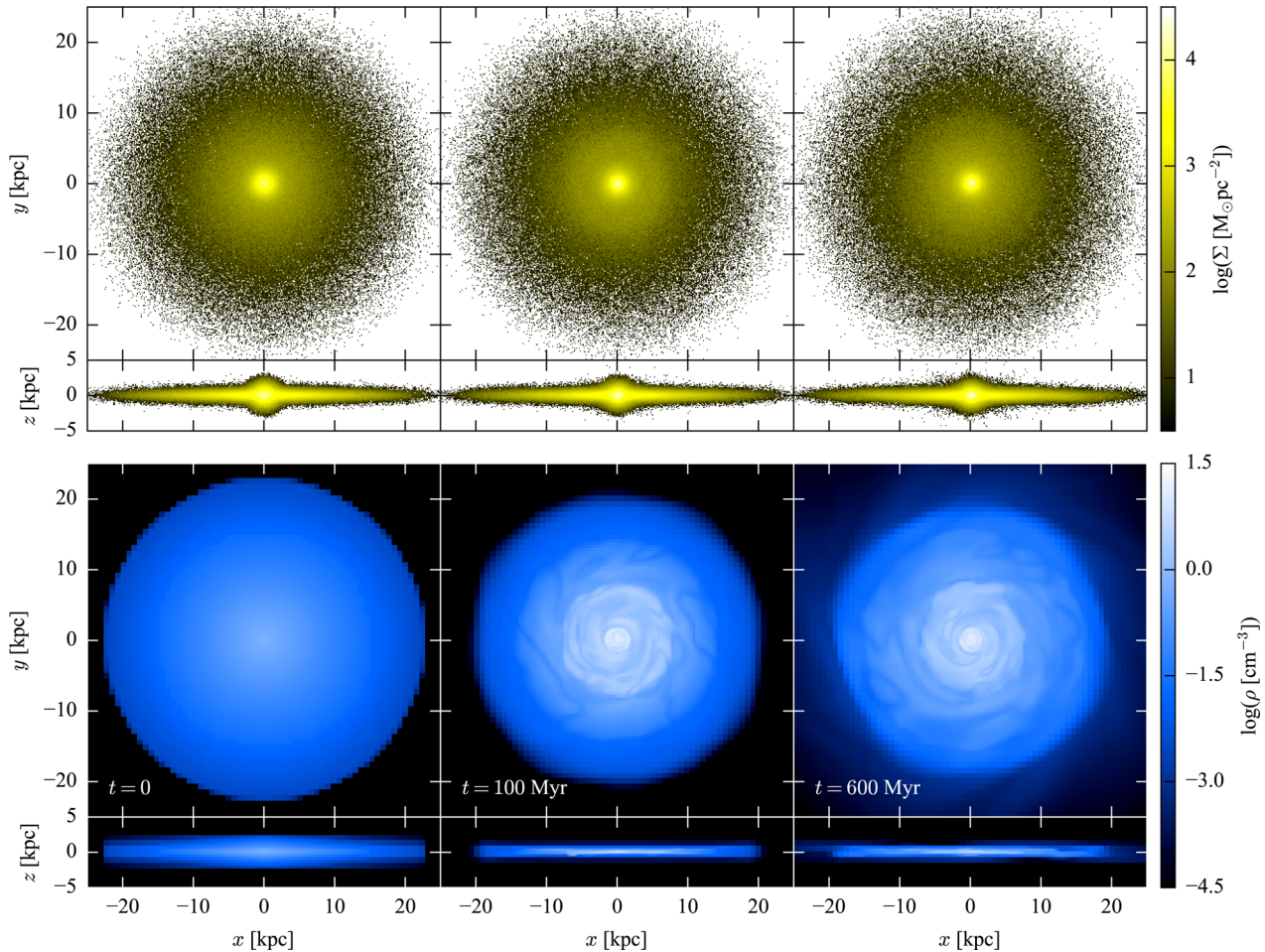


Figure B1. Face-on and edge-on surface density map of the stellar component (top) and gas volume density averaged along the line of sight (bottom) for the Newtonian model run in isolation. The left column represents the initial conditions. The central and right columns are 100 and 600 Myr later, showing the virialization from the initial conditions and the long-term evolution, respectively.

case and the DM in the Newtonian case. This figure represents the imperfections of our assumption consisting in replacing the DM with the MOND formalism. Except in the central kpc, the DM halo has a stronger gravitational influence on the baryons than the phantom DM, where differences can locally reach 30 per cent. In these regions, the stellar component dominates the mass budget of the galaxy (Fig. B3). The differences however become much smaller at larger radii, where the non-baryonic component takes over. Such differences account for the enhanced formation of substructures in the Milgromian model noted above. The differences remaining relatively small, and because of the simplicity of the method allowing for a fair comparison of the two frameworks, we use these models for our simulations of interacting galaxies.

In short, for both models, the overall mass budget remains remarkably constant over several rotation periods, although the disc of the Milgromian model is prone to the formation of substructures after a few rotation periods only.

APPENDIX C: EXTENDED HALOES

Truncating DM haloes represents a significant gain in memory and computational speed for simulations, and the model MW-A (Kuijken & Dubinski 1995) we use for our simulations takes

advantage of such simplification. However, in interacting galaxies, tidal tails extend to large distances and the gravitational potential they experience might be affected by the artificial truncation. In particular, Bournaud et al. (2003) show that the formation of TDGs requires extended DM halo to support the fragmentation of the tails.

To ensure that our results are not affected by the truncation artefact, we replace the lowered Evans halo with a more extended one, based on the NFW profile (Navarro et al. 1997). To do so, we fit the rotation curve of the naked DM halo of the MW-A model within the inner 30 kpc with an NFW profile (Fig. C1). The best-fitting NFW profile has a concentration of 13.7, a scale radius of 13.4 kpc and a ‘virial’ mass of $M_{200} \approx 6.65 \times 10^{11} M_{\odot}$, leading to a ‘virial’ radius of $r_{200} \approx 183$ kpc. Using this set of parameters, we realize a complete galaxy model, keeping the baryonic components unchanged⁸ and place it on the N_A orbit.⁹ Since the large-scale gas flows and small-scale cloud physics are not affected by the change of DM

⁸ For practical reasons, the new halo is truncated at 200 kpc.

⁹ The high mass of the new progenitor makes the interaction faster than in our fiducial model. To ensure that our results on the effects of the extended halo are not affected by this change of the orbit, we also ran models with reduced velocities, and reach the same conclusions.

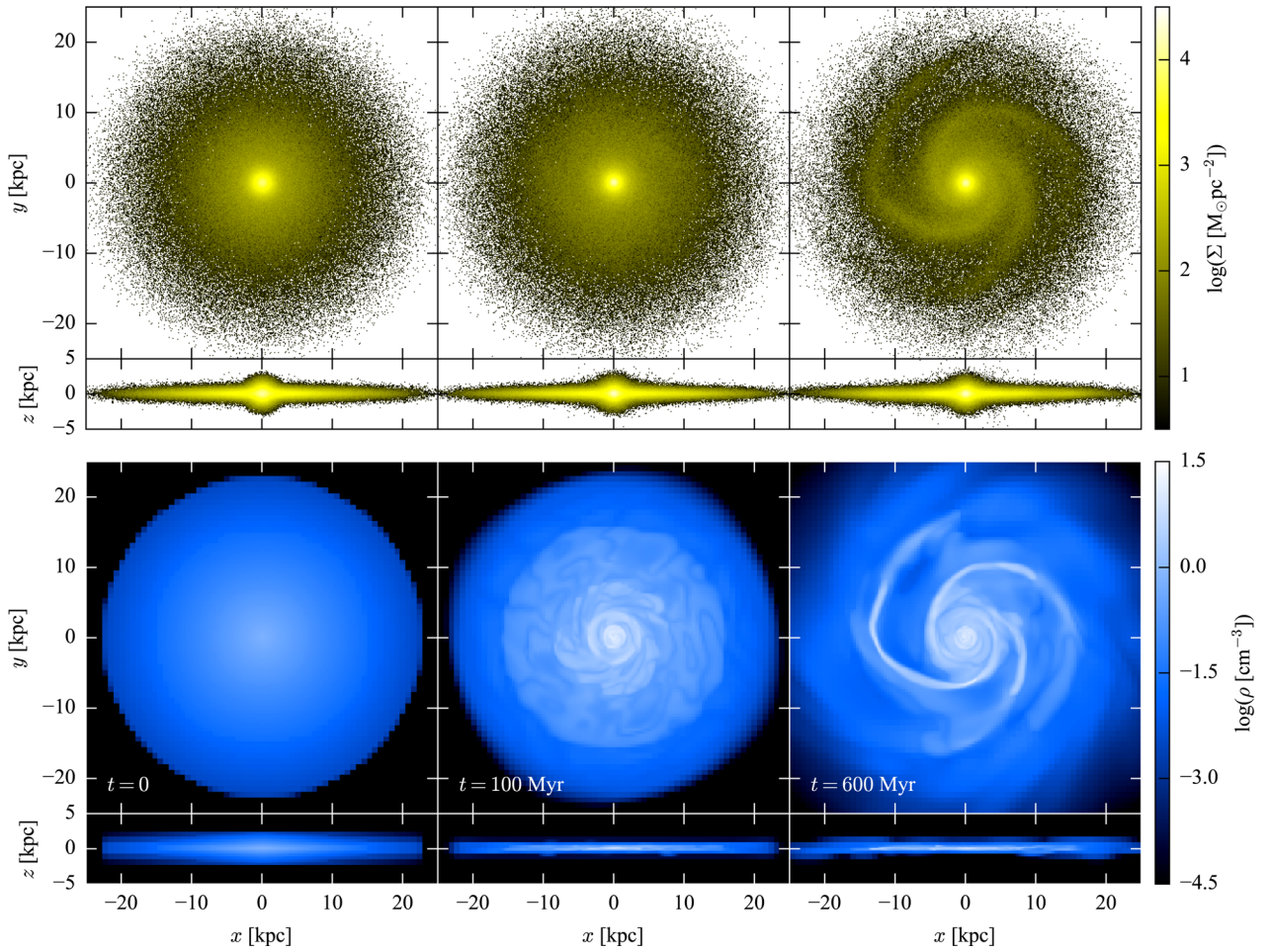


Figure B2. Same as Fig. B1, but for the Milgromian model. By construction, the initial conditions (left column) are strictly identical to the Newtonian case.

distribution in the outer parts of the halo, the SFR of this additional model is comparable within a few per cent to that of the truncated models. The main differences occur at coalescence, likely because of the slight differences in the orbit induced by the higher total mass of the progenitors.

Fig. C2 shows that no substructure nor agglomeration is seen in the stellar or gaseous tidal tails, in a similar way as our more severely truncated fiducial model. We conclude that the possible formation of substructures in the tidal debris is not affected by our choice of truncation of the DM halo of our galaxy progenitors.

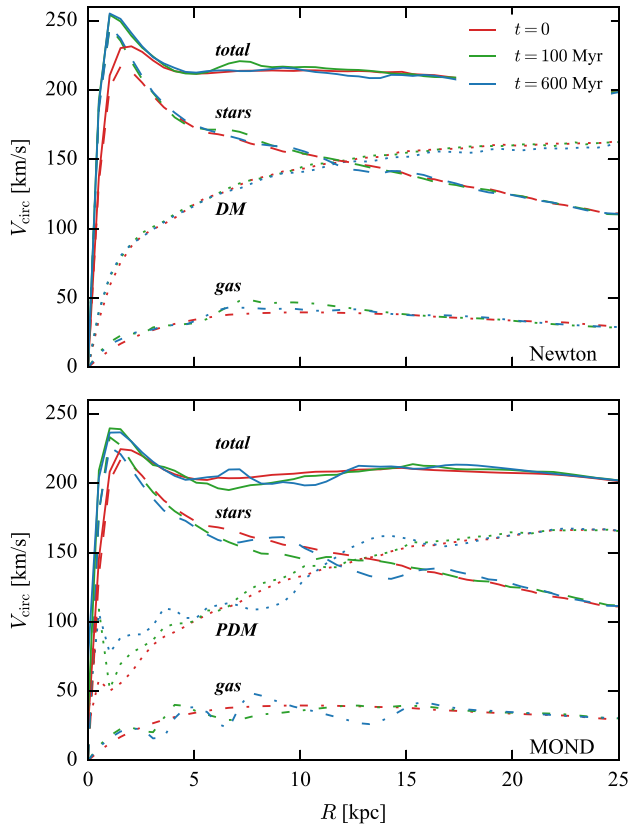


Figure B3. Rotation curves (solid lines) of our Newtonian (top) and Milgromian (bottom) models at the three instants shown in Figs B1 and B2. Dashed and dash-dotted lines represent the rotation curves for the stellar and gaseous components, respectively. The dotted lines correspond to the DM (Newton) and the phantom DM (MOND). The latter is computed by subtracting the baryonic contributions to the total circular velocity (assuming Newtonian physics), itself computed from the mid-plane total potential given by POR.

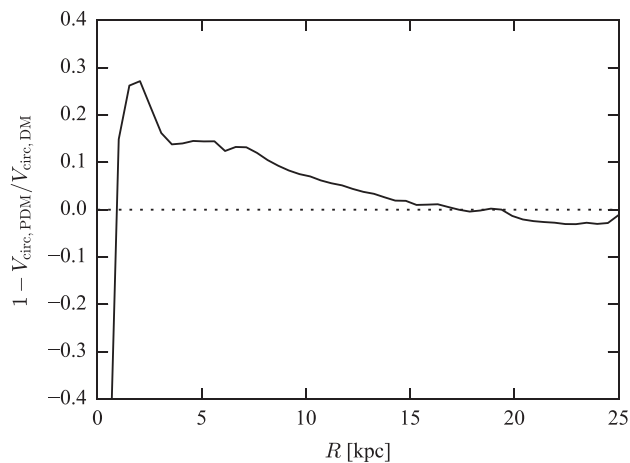


Figure B4. Relative difference between the circular velocity from the phantom DM and that from the DM, from our initial conditions.

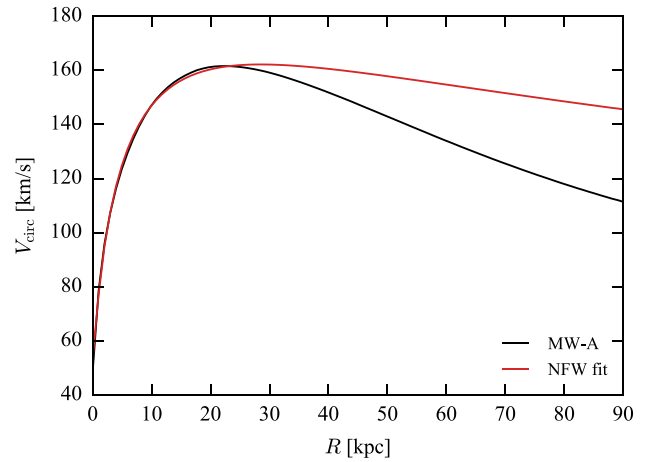


Figure C1. Circular velocity of the MW-A model (truncated at 50 kpc) used in this work, and a fit within the inner 30 kpc using a non-truncated NFW profile.

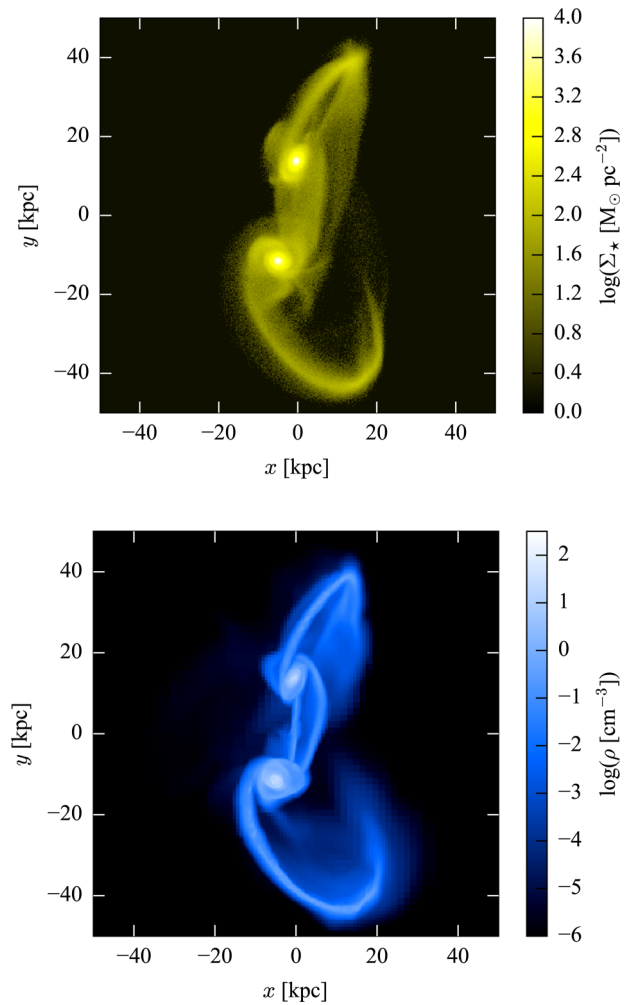


Figure C2. Map of the stellar surface density (top) and the gas volume density (bottom) of the interacting model with extended halo. No substructure is visible in the tidal tails.

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.