Persistence of the colour–density relation and efficient environmental quenching to \( z \sim 1.4 \)


\(^{1}\)Department of Physics, University of California, Davis, One Shields Ave., Davis, CA 95616, USA
\(^{2}\)Institute for Astronomy, University of Hawai‘i, 2680 Woodlawn Drive, Honolulu, HI 96822, USA
\(^{3}\)Max-Planck Institut für Astronomie, Königstuhl 17, D-69117 Heidelberg, Germany
\(^{4}\)National Astronomical Observatory of Japan, Osawa 2-21-1, Mitaka, Tokyo 181-8588, Japan
\(^{5}\)University of Paris Denis Diderot, University of Paris Sorbonne Cité (PSC), F-75205 Paris Cedex 13, France
\(^{6}\)Sorbonne Université, Observatoire de Paris, Université PSL, CNRS, LERMA, F-75014 Paris, France
\(^{7}\)Jet Propulsion Laboratory, Cahill Center for Astronomy & Astrophysics, California Institute of Technology, 4800 Oak Grove Drive, Pasadena, CA, USA
\(^{8}\)Aix Marseille Université, CNRS, LAM (Laboratoire d’Astrophysique de Marseille) UMR 7326, F-13388 Marseille, France
\(^{9}\)National Center for Supercomputing Applications, University of Illinois, 1205 West Clark St., Urbana, IL 61801, USA
\(^{10}\)Department of Physics and Astronomy, Colby College, Waterville, ME 04961, USA
\(^{11}\)Spitzer Science Center, California Institute of Technology, M/S 220-6, 1200 E. California Blvd., Pasadena, CA 91125, USA

Accepted 2019 September 9. Received 2019 August 30; in original form 2018 November 28

ABSTRACT

Using \( \sim 5000 \) spectroscopically confirmed galaxies drawn from the Observations of Redshift Evolution in Large Scale Environments (ORELSE) survey we investigate the relationship between colour and galaxy density for galaxy populations of various stellar masses in the redshift range \( 0.55 \leq z \leq 1.4 \). The fraction of galaxies with colours consistent with no ongoing star formation (\( f_q \)) is broadly observed to increase with increasing stellar mass, increasing galaxy density, and decreasing redshift, with clear differences observed in \( f_q \) between field and group/cluster galaxies at the highest redshifts studied. We use a semi-empirical model to generate a suite of mock group/cluster galaxies unaffected by environmentally specific processes and compare these galaxies at fixed stellar mass and redshift to observed populations to constrain the efficiency of environmentally driven quenching (\( \Psi_{\text{convert}} \)). High-density environments from \( 0.55 \leq z \leq 1.4 \) appear capable of efficiently quenching galaxies with \( \log(\mathcal{M}_*/M_\odot) > 10.45 \). Lower stellar mass galaxies also appear efficiently quenched at the lowest redshifts studied here, but this quenching efficiency is seen to drop precipitously with increasing redshift. Quenching efficiencies, combined with simulated group/cluster accretion histories and results on the star formation rate-density relation from a companion ORELSE study, are used to constrain the average time from group/cluster accretion to quiescence and the elapsed time between accretion and the quenching event. These time-scales were constrained to be \( \langle t_{\text{convert}} \rangle = 2.4 \pm 0.3 \) and \( \langle t_{\text{delay}} \rangle = 1.3 \pm 0.4 \) Gyr, respectively, for galaxies with \( \log(\mathcal{M}_*/M_\odot) > 10.45 \) and \( \langle t_{\text{convert}} \rangle = 3.3 \pm 0.3 \) and \( \langle t_{\text{delay}} \rangle = 2.2 \pm 0.4 \) Gyr for lower stellar mass galaxies. These quenching efficiencies and associated time-scales are used to rule out certain environmental mechanisms as being the primary processes responsible for transforming the star formation properties of galaxies over this 4 Gyr window in cosmic time.

Key words: techniques: photometric – techniques: spectroscopic – galaxies: clusters: general – galaxies: evolution – galaxies: groups: general.
1 INTRODUCTION

One of the most fundamental properties associated with the study of galaxy evolution is whether or not a galaxy is forming stars. This information, when determined for a large and diverse enough population of galaxies, can be used to directly test and to calibrate simulations and models of galaxy evolution (e.g. McGee et al. 2011; Tal et al. 2014; Balogh et al. 2016; Kawinwanichakij et al. 2016, 2017; Socolovsky et al. 2018). While the processes governing the likelihood that a particular galaxy is actively forming stars are extremely complex, and, further, while stochastic rather than smooth star formation histories may be typical for galaxies (e.g. Pacifici et al. 2013; Teyssier et al. 2013; Kauffmann 2014; Sparre et al. 2017; Emami et al. 2019), when averaging over an ensemble of galaxies there are physical parameters that can be used as strong predictors of how many galaxies in that population are likely to be forming stars. At least to $z \sim 1$, it appears that three parameters in particular, cosmic epoch, the mass of the stellar content contained within the galaxy, and the environment in which a galaxy resides, correlate well with the star formation properties of galaxies. While there are many other properties beyond these three which may serve to modulate star formation including, but not limited to, the level of active galactic nucleus (AGN) activity (e.g. Kocevski et al. 2009; Rumbaugh et al. 2012; Dubois et al. 2013; Juneau et al. 2013; Lemaux et al. 2014b), shocks, merging, and other galaxy-wide processes (e.g. Birnboim & Dekel 2003; Kereš et al. 2005; López-Sanjuan et al. 2013; Belfiore et al. 2016, 2017), kinematics (e.g. Genzel et al. 2008, 2014; Hung et al. 2019), morphology (e.g. Martig et al. 2009; Bell et al. 2012; Whitaker et al. 2017), and the amount and metallicity of the gas, if any, contained within galaxies (e.g. Scoville et al. 2016, 2017; Darvish et al. 2018; Sanders et al. 2018; Wang et al. 2018), all of these properties can be thought of in some way to be subordinate (i.e. correlated with) to redshift, stellar mass, or environment or some combination of the three.

However, while studies exist in abundance probing each of these three parameters in tandem across a wide variety of redshifts, stellar masses, and environments $z \sim 1$, there are few studies which simultaneously and self-consistently probe a wide baseline in each of the three parameters. For example, Wetzel et al. (2013) presented an extensive investigation of the star formation rate (hereafter $\text{SFR}$) distribution and the fraction of galaxies not forming stars, otherwise known as the quiescent fraction ($q_{\text{q}}$), for a large sample of spectroscopically confirmed galaxies from the Sloan Digital Sky Survey (SDSS). While this study, by virtue of the large number of galaxies, and the large dynamic range of stellar masses and environments probed, provided valuable insights on the star formation histories of galaxies in different environments, these insights were broadly limited to galaxy populations residing in the local universe ($z \lesssim 0.3$). At higher redshift ($z \sim 1$), samples from the VIMOS Very Deep Survey (VVDS; Le Févre et al. 2005, 2013), the Deep Extragalactic Evolutionary Probe 2 (DEEP2; Davis et al. 2003; Newman et al. 2013), the zCOSMOS-Bright survey (Lilly et al. 2007, 2009), and the VIMOS Public Extragalactic Redshift Survey (VIPERS; Garilli et al. 2014; Guzzo et al. 2014) have been extensively studied to measure variations of $q_{\text{q}}$. In such surveys, this quantity is typically proxied by the number of red galaxies relative to the total number of galaxies, as a function of redshift, stellar mass (or, somewhat equivalently, absolute magnitude), and environment (e.g. Cucciati et al. 2006; Cooper et al. 2007; Peng et al. 2010; Kováč et al. 2014; Cucciati et al. 2016). However, the range of environments probed by each survey is typically limited to, at most, massive group environments, and the spectroscopic populations are typically skewed toward bluer galaxies. As such, the conclusions drawn on these results necessarily share the same limitations. Conversely, observational campaigns exist which target the cores and intermediate regions of massive clusters at a variety of redshifts, but such campaigns can suffer from a spectroscopic sampling issue, lack of a comparable field sample, large sample variance from the small number of clusters probed, small redshift windows, and the lack of sampling in the sparser regions of clusters and the surrounding filamentary structure.

While the number of conflicting results and the magnitude of these conflicts speak to the difficulty of overcoming the above difficulties completely, clever methods have been employed to overcome some of these limitations, and such campaigns have been used to good effect to present compelling evidence that the relationship of the quiescent fraction of galaxies with environment, stellar mass, and redshift persists to some extent up to $z \sim 1.5$ (e.g. Strazzullo et al. 2010; Snyder et al. 2011; Quadri et al. 2012; Cooke et al. 2016; Nantais et al. 2016). Secular processes strongly tied to the stellar mass content of galaxies are generally thought to regulate star formation in lower-density environments, with star-forming galaxies observed to be less prevalent amongst galaxies at higher stellar mass. In these rarefied environments, small environmental effects may enter for galaxies that reside in or near voids and filaments, as the distance to the void or filament appears to modulate the colours and star-forming properties of such galaxies (e.g. Moorman et al. 2016; Kuutma, Tamm & Tempel 2017; Kraljic et al. 2018; Laigle et al. 2018), however, it generally appears that secular processes dominate. While the trend of increasing quiescent fraction with the increasing mass of stellar content in field environments appears to hold to $z \sim 2$–3 (e.g. Ilbert et al. 2013; Muzzin et al. 2013; Tomczak et al. 2014), the quiescent fraction at which quiescent galaxies become more abundant than star-forming galaxies is a strong function of redshift. Another metric used to probe the star-forming state of galaxies is the star formation rate normalized to the stellar mass content, otherwise known as the specific star formation rate ($\text{SSFR}$). This metric appears, both for all types of galaxies and for star-forming galaxies only, to be strongly tied to both the redshift and stellar mass, in that the average $\text{SSFR}$ of galaxies is seen to increase at increasing redshift and for galaxy populations with smaller average stellar masses (e.g. Tomczak et al. 2016).

When a larger dynamic range of environments is probed, including the intermediate- to high-density regimes of groups and clusters, the situation becomes more complex. While, among star-forming galaxies, at fixed $M_*$ the average $\text{SSFR}$ appears to be largely independent across all types of environments (e.g. Muzzin et al. 2012; Wijesinghe et al. 2012; Koyama et al. 2013; Zeimann et al. 2013; Darvish et al. 2016; Vulcani et al. 2017; Wagner et al. 2017, though see also Patel et al. 2011; Ziparo et al. 2014; Tran et al. 2015; Tomczak et al. 2019), the quiescent fraction is consistently found to be a strong function of environment at least until $z \sim 1$ (e.g. Balogh et al. 2016; Cucciati et al. 2017; Kawinwanichakij et al. 2017; Jian et al. 2018). This behaviour indicates that the ability of galaxies to form stars is in some way related to the plethora of processes which are enhanced or exclusive to intermediate- to high-density environments. Indeed, the behaviour of the quiescent fraction as a function of stellar mass and environment at fixed redshift has been used in a number of studies to estimate time-scales associated with the quenching of group and cluster galaxies as well as the processes involved in that quenching (e.g. Mok et al. 2014; Balogh et al. 2016; Fossati et al. 2017; Foltz et al. 2018). While promising, the limited number of cluster and group environments studied, the
limited number of galaxies studied per system, and the lack of a comprehensive investigation of quenching time-scales as a function of both stellar mass and redshift, make it difficult to understand how applicable the conclusions drawn from these investigations are to the evolution of generic galaxy populations in higher density environments.

In this paper we present a study of a large sample (~5000) of spectroscopically confirmed galaxies from the Observations of Redshift Evolution in Large Scale Environments (ORELSE; Lubin et al. 2009) survey, an immense imaging and spectroscopic campaign over ~5° probing the surrounding structure of 15 known large-scale structures at 0.6 ≤ z ≤ 1.3 on ~10–15 h−1 Mpc scales. For this sample, which is largely representative of the overall galaxy population for galaxies with log(M*/M⊙) > 10, we measure f′r at different redshifts, stellar masses, and environments. These measurements, in conjunction with semi-empirical models and N-body simulations, are used to self-consistently infer the efficiency with which group and cluster environments are able to quench star formation in their constituent galaxies to z ∼ 1.4 and the quenching time-scale for galaxies of different stellar masses. In tandem with a companion study (Tomczak et al. 2019) in which a similar ORELSE sample is used to study the relationship between SSFR and environment for star-forming galaxies, we further estimate the time-scale that galaxies of differing stellar masses can remain in group and cluster environments without undergoing a galaxy-scale quenching event. These time-scales are subsequently used to argue in favour or against certain processes as agents of this quenching.

Throughout this paper all magnitudes, including those in the IR, are presented in the AB system (Oke & Gunn 1983; Fukugita et al. 1996). All distances are quoted in proper units. We adopt a concordance Λ cold dark matter cosmology with H0 = 70 km s−1 Mpc−1, ΩΛ = 0.73, and ΩM = 0.27. While abbreviated throughout the paper for convenience, stellar masses are presented in units of M70 h−2 M⊙, star formation rates in units of h−2 M⊙ yr−1, total and halo masses in units of h−1 M⊙, ages in units of h−1 Gyr or h−1 Myr, absolute magnitudes in units of MAB + 5 log(h70), proper distances in units of h−1 kpc Mpc−1, where h70 = H0/70 km s−1 Mpc.  

2 THE ORELSE FIELD AND LARGE-SCALE STRUCTURE SAMPLE

In this paper we incorporate samples drawn from observations of all 15 of the final ORELSE fields. Adopting the naming convention of Lubin et al. (2009), these fields are Cl 0023+0423, RCS J0224−0002, XLSSC 005a+b, Cl J0849+4452, Cl J0910+5422, RX J1053.7+5735, 1137+3000, RX J1121.4+4918, Cl J1324+3011, Cl J1324+3059, Cl J1350+6007, Cl J1429.0+4221, Cl J1604+4304/Cl J1604+4321, RX J1716.4+6708, RX J1745.2+6556. The full galaxy sample taken from these fields is comprised of nearly 9000 spectroscopically confirmed galaxies in the redshift range studied in this paper, 0.55 ≤ z ≤ 1.4, which span a range of environments from the sparsely populated field to the cores of massive, relaxed clusters. In total, 56 groups and clusters were discovered across all 15 of the ORELSE fields spanning the ranges of redshift, line-of-sight (LOS) galaxy velocity dispersion, and virial masses shown in the three panels of Fig. 1. The latter of these three measurements are made using the relation given in Lemaux et al. (2012). While several thousand of the galaxies in the sample presented here are confirmed members of either these groups and clusters or the large-scale structures (LSSs) that they are embedded in, the remainder of our sample, consisting of several thousand more galaxies, are completely unassociated with any massive structure. Such samples allow for an internal comparison of various subsets of galaxies across a broad range of environments at different epochs. Imaging and spectroscopic observations made on nearly all of the final 15 ORELSE fields, as well as methods relating to the reduction of these data and the measuring of relevant quantities from these observations, have been discussed in other ORELSE studies (with the exception of one field, see Section 2.3). As such, we only briefly describe the characteristics of these observations and methods here.

2.1 Imaging and photometry

Optical r′z′ imaging for most ORELSE fields was initially performed using either the Large Format Camera (LFC; Simcoe et al. 2000) situated on the prime focus of the Palomar 5-m Hale telescope or Suprime-Cam (Miyazaki et al. 2002) mounted on the prime focus of the Subaru 8-m telescope. In one case, XLSSC005a+b (hereafter XLSS005), initial optical imaging was drawn from Megacam (Boulade et al. 2003) observations from the 3.6-m Canada–France–Hawaii Telescope (CFHT) taken as part of the ‘Deep’ portion of the CFHT Legacy Survey (CFHTLS). Additional B and V imaging was taken for all ORELSE fields with Subaru/Suprime-Cam except for the XLSS005 field where r′- and g-band imaging were available. In some cases Subaru/Suprime-cam imaging in the Y band was also taken. The typical depths of the optical imaging ranged from mAB = 26.4 in the B band to mAB = 24.6 in the z′/I′ bands as estimated from the method described in Tomczak et al. (2017). All LFC data were reduced with Image Reduction and Analysis Facility (IRAF; Tody 1993) following the methods of Gal et al. (2008). Reduction of the Suprime-Cam data were performed with the SDFRED2 pipeline (Ouchi et al. 2004) supplemented by several Traitement Élémentaire Rédaction et Analyse des PIlXes (TERAPIX) routines2 (for more details, see Tomczak et al. 2017). In both cases, photometric calibration was performed from observations of Landolt (1992) standard star fields taken on the same night of each observation. CFHTLS observations were reduced and photometrically calibrated using TERAPIX routines following the methods described in Ilbert et al. (2006) and the T0006 CFHTLS handbook. Though ORELSE imaging observations in the redder portion of the observed-frame optical employed a variety of different filter sets, for simplicity, we will usually use the generalized terms r, i, and z throughout the paper to refer to all variants of these filters (i.e. r′, i′, and z′, respectively) except when specificity is helpful.

All but one field in ORELSE (Cl1350) was observed in the near-infrared (NIR) J and K/K′ bands. These observations were taken from the Wide-Field Camera (WFCAM; Casali et al. 2007) mounted on the United Kingdom Infrared Telescope (UKIRT) and the Wide-field InfraRed Camera (WIRCam; Puget et al. 2004) mounted on CFHT and reached a typical depth of mAB = 21.9 and 21.7 for the J

1http://terapix.iap.fr

2http://terapix.iap.fr/cplt/T0006-doc.pdf

3C. D. Fassnacht, private communication.

4http://terapix.iap.fr

Downloaded from https://academic.oup.com/mnras/article/490/1/1231/5573279 by guest on 27 January 2021
and K/Ks bands, respectively. The UKIRT data were processed using the standard UKIRT processing pipeline provided courtesy of the Cambridge Astronomy Survey Unit\(^3\) and, the CFHT data, through the I’iwi pre-processing routines and TERAPIX. Both pipelines provided fully reduced mosaics and associated weight maps. The photometric calibration of the mosaics output by both pipelines was done selecting bright (\(m < 15\)), non-saturated objects with existing Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006) photometry, with appropriate \(k\)-corrections made using stars drawn from the Infrared Telescope Facility (IRTF) spectral library (Rayner, Cushing & Vacca 2009). Additional imaging in the NIR was taken from the Spitzer (Werner et al. 2004) space observatory using the InfraRed Array Camera (IRAC; Fazio et al. 2004) in the two non-cryogenic channels (\([3.6]/[4.5]\)) for all 15 ORELSE fields and additionally in the two cryogenic channels (\([5.8]/[8.0]\)) for four of the ORELSE fields (SC1604, RXJ1716, RXJ1053, and XLSS005 following the naming convention of Rumbaugh et al. 2018; Tomczak et al. 2019) to an average depth of 24.0, 23.8, 22.4, and 22.3 mag, respectively. The basic calibrated data (cBCD) images provided by the Spitzer Heritage Archive were reduced using the MOsaicker and Point source EXtractor (MOPEX; Makovoz & Marleau 2005) package in conjunction with several custom Interactive Data Language (IDL) scripts written by J. Surace. For more details on the reduction of this data, see Tomczak et al. (2017).

For each field, all optical and non-Spitzer images were registered to a common grid of plate scale 0.2 arcsec pixel\(^{-1}\) and convolved to the worst point spread function (PSF) for that field using the methods described in Tomczak et al. (2017). Prior to this convolution, some images with exceptionally large PSFs for which we had the luxury of another image with broadly redundant spectral coverage (i.e. LFC \(r'\) versus Suprime-cam \(R'\) imaging) with a smaller measured PSF were removed. The worst PSF per field for all retained ground-based optical/NIR images ranged from \(\sim 1.00\) to 1.96 arcsec for the fields studied here, with only one field (CII1350) having an image with a PSF that exceeded 1.4 arcsec. Source detection and photometry were obtained by running Source Extractor (SEXTRACTOR; Bertin & Arnouts 1996) in dual-image mode using either a stacked \(\chi^2\) optical image or an image in a single band as a detection image (for details on the specific image used for each field except for 1137+3000 (hereafter CII1137), see Tomczak et al. (2017) and Rumbaugh et al. (2018). For details on CII1137, see Section 2.3. Fixed-aperture photometry was performed on all PSF-matched images with SEXTRACTOR employing an aperture of \(1.3\times\) the FWHM of the homogenized PSF for each field and transformed to a total magnitude using the ratio of aperture and AUTO flux densities as measured in the detection image. Magnitude uncertainties were calculated from adding, in quadrature, SEXTRACTOR uncertainties to our own estimates of background noise drawn from the 1\(\sigma\) root mean square (RMS) scatter of measurements in hundreds of blank sky regions for each band. Spitzer/IRAC magnitudes were incorporated by running the software T-PHOT (Merlin et al. 2015) on the fully reduced mosaics using the segmentation maps from the ground-based detection images as input, with flux density uncertainties estimated from the scaled best-fitting model for each object. For more details on the reduction and measurements of ORELSE imaging data, see Tomczak et al. (2017).

### 2.2 Optical spectroscopy

The optical imaging described above was used following the methods described in Lubin et al. (2009) to select targets for spectroscopic follow-up with the DEep Imaging and Multi-Object Spectrometer (DEIMOS; Faber et al. 2003), located at the Nasmyth focus of the Keck II telescope. In brief, the spectroscopic targeting scheme employed a series of colour and magnitude cuts that are applied to maximize the number of targets with a high likelihood of being on the cluster/group red sequence at the presumed redshift of the LSS in each field (i.e. priority 1 targets). However, as discussed extensively in Tomczak et al. (2017), because of the relative rarity of such objects, the majority of targets for all ORELSE fields were objects with colours outside of these ranges (predominantly blueward). The fraction of priority 1 targets which entered into our final sample ranged from \(\sim 10\) per cent to \(\sim 45\) per cent across all ORELSE fields, a fraction which tended to vary strongly with

\(3\)http://casu.ast.cam.ac.uk/surveys-projects/wfcam/technical

---

**Figure 1.** Left: Redshift distribution of the 56 groups and clusters detected in the 15 ORELSE fields presented in this paper. The systemic redshift of each structure is estimated by the average of all spectral members within 1 Mpc of the luminosity-weighted spectral member centre (see Ascaso et al. 2014). Solid and dashed vertical lines denote the median redshift of ORELSE clusters \((\sigma_{z,\text{LOS}} \geq 550 \text{ km s}^{-1})\) and groups \((\sigma_{z,\text{LOS}} < 550 \text{ km s}^{-1})\), respectively. Centre: Distribution of \(\sigma_{z}\) for the same 56 structures. Velocity dispersions are calculated from an average of 34 and 15 member galaxies for ORELSE clusters and groups, respectively. The median \(\sigma_{z}\) of clusters and groups are indicated by the solid and dashed vertical lines, respectively. The average \(\sigma_{\text{LOS}}\) error, estimated through jackknife resampling for each structure, is shown in the top left. Right: Virial mass of all 56 structures as estimated from \(\sigma_{v}\). The median \(M_{\text{vir}}\) for clusters and groups are indicated by solid and dashed vertical lines, respectively. The average \(\Delta M_{\text{vir}}\) error is shown in the top left.
the density of spectroscopic sampling per field. We discuss the consequences of this targeting scheme for the results presented in this study in Section 2.5. Spectroscopic targets were generally limited to $i < 24.5$, with a median of $i \sim 23$, though this magnitude limit was not imposed strictly and an appreciable number of objects were targeted below this limit. In addition, objects detected at X-ray wavelengths in our Chandra imaging (Rumbaugh et al. 2017) or at radio wavelengths from our Very Large Array (VLA) 1.4 GHz imaging (Shen et al. 2017) were also highly prioritized, though the number density of these objects is typically low and, thus, such objects only comprised a small fraction of targets on a given mask (typically $\lesssim 10$ per cent).

Spectral observations with DEIMOS taken as part of ORELSE exclusively used the 1200 l mm$^{-1}$ grating with 1 arcsec slit widths for the $\sim 120$ targets per slitmask and a central wavelength that ranged between 7000 and 8700 Å depending on the redshift of the LSS being targeted. These observations resulted in a pixel scale of 0.33 Å pixel$^{-1}$, a wavelength coverage of $\pm 1300$ Å roughly centred on the central wavelength of the observation, and a spectral resolution of $R \sim 5000 (\lambda/\delta \lambda_{FWHM})$, where $\delta \lambda_{FWHM}$ is the FWHM spectral resolution). This set-up and resolution was required to allow us the possibility to detect several important spectral features at the redshift of the targeted LSS (e.g. [O II] $\lambda 3726, 3729$ Å, Ca II K & $\lambda 3934, 3969$ Å, D$_{\text{h}}$($4000$), and H $\&$ $\lambda 4101$ Å) and to ensure, in most cases, the spectral separation of the $\lambda 3726, 3729$ Å [O II] doublet. Separating this doublet allows for a secure determination of the redshift of targets based on this feature alone. The number of slitmasks per field varied from 4 (RC0224) to 18 (SC1604), with generally larger and more complex LSSs, as well as those at higher redshift, given more extensive coverage. Average integration times per mask ranged from $\sim 7000$ to $\sim 10 500$ s and were varied based on conditions and the faintness of the target population to roughly maintain the same median signal-to-noise ratio of all targets from mask to mask. For details on the observations of specific ORELSE fields, see Table 1.

Observations from DEIMOS were reduced using a modified version of the Deep Evolutionary Extragalactic Probe 2 (DEEP2; Davis et al. 2003; Newman et al. 2013) SPEC2d pipeline. This package combines the individual exposures of the slit mosaic and performs wavelength calibration, cosmic ray removal and sky subtraction on a slit by slit basis, generating a processed two-dimensional spectrum for each slit. The SPEC2d pipeline also generates a processed one-dimensional spectrum for each slit. This extraction creates a one-dimensional spectrum of the target, containing the summed flux at each wavelength in both a boxcar and an optimized window (Horne 1986). The accompanying SPEC1d package is then run on all resulting one-dimensional spectra. This package cross-correlates a suite of galactic and stellar templates to find 10 redshifts which correspond local minima in $\chi^2$ space for different combinations of templates. These redshifts are used to inform the visual inspection process performed subsequent to the reduction and cross-correlation steps (see below).

For the purposes of the reduction of the data presented in this paper, several modifications were made to the official version of the SPEC2d pipeline. These modifications included an improved methodology of interpolating over the $\sim 10$ Å chip gap between the red and blue CCD arrays on DEIMOS, the implementation of an improved throughput correction, and an improvement in the functionality related to estimating the final wavelength solution. More specifically, rather than interpolating over single pixel values on either side of the gap, the pipeline now smooths variations over relatively large windows ($\sim 50$ pixels) on either side of the chip gap and linearly interpolates over these smoothed arrays. This implementation helps considerably in reducing artefacts resulting from reduction or from improper identification of the gap. In addition, an improved method of applying the measured instrumental throughput was employed by combining throughput data from multiple discrete observational set-ups$^7$ and estimating an appropriate throughput correction at every wavelength for observational set-ups at all filter/grating/tilt combinations. The pipeline also now attempts fits to multiple optical models if the fiducial guess fails, has an increased search window for tweaking the wavelength solution resulting from the initial optical model guess as well as an increased tolerance for offsets between wavelength solutions derived from arc lamps and those derived from night sky lines. Additionally, a more complete night sky line list has been added to the code to allow for more accurate matching to the airglow spectrum measured on the science frames. These latter changes were necessary as it was noticed early on in the reduction of some masks that wavelength solutions were either failing catastrophically, or, worse, converging, without error, to incorrect values for some slits on our masks while converging to correct values for other slits. Finally, a larger suite of empirical templates generated from a variety of observations (Lilly et al. 2007; Lemaux et al. 2009; Le Fèvre et al. 2013, 2015) was added to the templates used in the SPEC1d cross-correlation process primarily to allow the extension of its functionality to higher ($z > 2$) redshift.

Following the reduction and template cross-correlation, each spectrum was then visually inspected using the publicly available DEEP2 redshift measurement program, ZSPEC (Newman et al. 2013) to determine, if possible, the redshift of each target either from the list 10 redshifts determined by SPEC1d or from a custom redshift fit to visually identified spectral features. Additionally, all two-dimensional spectra were searched for serendipitous detections both spatially coincident and separated from target galaxies (see Lemaux et al. 2009 for details on these types of detections and the method used for finding them). For those spectra which contain one or more serendipitous detections, one-dimensional spectra were extracted in the same manner as was done with the SPEC2d pipeline and their redshifts were determined, when possible, in the zspec environment. Each target and serendipitous detection was assigned a quality code, $Q$, which represents our confidence in the redshift measurement. The criteria required to assign each quality code is the same as that adopted in the DEEP2 survey (see Gal et al. 2008; Newman et al. 2013; Tomczak et al. 2017 for details on these criteria). For this study, we consider only those objects with $Q = -1, 3,$ and 4 to have secure spectral redshifts. These three quality codes correspond to stellar ($Q = -1$) and extragalactic ($Q = 3, 4$) redshifts secure at the $>95$ per cent level.

Additionally, a small number of spectral redshifts ($\sim 250$) were included from ORELSE precursor surveys or other, unrelated surveys designed specifically to observe the LSSs also targeted in ORELSE (Oke, Postman & Lubin 1998; Gal & Lubin 2004; Gioia et al. 2004; Tanaka et al. 2008; Mei et al. 2012) that employed a variety of different instruments and set-ups. For spectra coming from these surveys we imposed, to the best of our ability, the same quality code system as was applied to our DEIMOS data and accepted only

---

$^6$The true central wavelength of a given slit can vary $\pm 150$ Å from the fiducial value depending on where the slit is placed on the DEIMOS mask in the direction parallel to the dispersion dimension.

$^7$Taken from https://www2.keck.hawaii.edu/inst/deimos/ripisc.html
Table 1. ORELSE DEIMOS spectral observations.a

| Field   | λc (Å) | Avg. coverage (Å) | t_int (s) | Avg. seeing (arcsec) | N_{mask} | N_{target} | N_{spec} | N_{spec-gal} | log(\mathcal{M}_{\text{vir}}/\mathcal{M}(\odot)^r | z_{spec} |
|---------|-------|------------------|----------|---------------------|---------|-----------|----------|-------------|------------------|
| SG0023  | 7500–7850 | 6200–9150         | 5700–9407 | 0.45–0.81          | 9       | 1128      | 923      | 758         | 12.7–13.9         | 0.829–0.980     |
| RSC0224 | 7300–7450 | 6000–8750         | 6840–7520 | 0.53–0.88          | 4       | 598       | 493      | 381         | 13.9–14.8         | 0.778–0.854     |
| XLSS005 | 7900     | 6600–9200         | 1063–12600 | 0.39–1.12         | 9       | 999       | 733      | 586         | 14.5             | 1.056           |
| SC0849  | 8700     | 7400–10000        | 6300–16200 | 0.51–1.50         | 8       | 977       | 556      | 390         | 12.7–14.7         | 0.568–1.270     |
| RXJ0910 | 8000–8100 | 6700–9400         | 7200–11664 | 0.50–1.05         | 5       | 971       | 736      | 505         | 12.7–14.7         | 0.760–1.103     |
| RXJ1053 | 8200     | 6900–9500         | 7200–9000 | 0.56–0.79         | 6       | 698       | 405      | 300         | 14.8             | 1.129–1.204     |
| CI1137  | 7850     | 6550–9150         | 2600–14400 | 0.53–1.20         | 6       | 827       | 539      | 438         | 14.1             | 0.955           |
| RXJ1221 | 7200     | 5900–8500         | 4800–8400 | 0.55–1.20         | 5       | 663       | 519      | 411         | 13.9–14.7         | 0.700–0.702     |
| SC1324  | 7200     | 5900–8500         | 2700–10800 | 0.44–1.00         | 12      | 1653      | 1332     | 985         | 12.8–14.8         | 0.696–1.098     |
| CI1350  | 7500     | 6200–8800         | 3600–10400 | 0.50–1.55         | 6       | 803       | 623      | 352         | 13.4–14.7         | 0.800–0.802     |
| CI1429  | 7400–7500 | 6100–8800         | 5000–7800 | 0.43–0.85         | 8       | 1017      | 835      | 563         | 14.8             | 0.987           |
| SC1604  | 7700     | 6400–9000         | 3600–14400 | 0.50–1.30         | 18      | 2358      | 1801     | 1294        | 13.4–14.7         | 0.600–1.182     |
| RXJ1716 | 7800     | 6500–9100         | 5400–9600 | 0.54–0.83         | 6       | 944       | 675      | 513         | 14.4–15.1         | 0.809–0.853     |
| RXJ1757 | 7000–7100 | 5700–8400         | 6300–14730 | 0.47–0.82         | 6       | 945       | 742      | 397         | 13.4–14.8         | 0.693–0.946     |
| RXJ1821 | 7500     | 6200–8800         | 7200–9000 | 0.58–0.86         | 6       | 728       | 611      | 342         | 14.5–15.1         | 0.817–0.919     |

aField names are shorthand notation of the names given in Section 2 and can be matched by number.

For numbers of targets, secure spectral redshifts, and secure spectral extragalactic redshifts, we also include numbers from all other spectral surveys incorporated into this study (see Section 2.2) with the exception of those from the VVDS survey.

Includes only those objects with a secure spectral redshift (see Section 2.2) and includes serendipitous detections which comprise ~3.6 per cent of the entire sample.

These numbers include only those galaxies with a secure spectral redshift in the redshift range used in this study, 0.55 ≤ z ≤ 1.4.

Viral mass range of groups and clusters detected within each field. These groups and clusters only refer to those previously known groups and clusters reported in Hung et al. (2019) and not those additional overdensities found in that study.

those spectral redshifts with a high probability of being correct (i.e. the equivalent of Q = −1, 3, and 4). Finally, an additional ~1000 redshifts were drawn from the VIMOS Very Deep Survey (VVDS; Le Fèvre et al. 2013), of which ~700 are within the redshift range used in this study (0.55 ≤ z ≤ 1.4). These objects are exclusively located in the XLSS005 field. For these data, we incorporated only those objects with spectroscopic reliability flags of X2, X3, X4, or X9, where X = 0–2,8 which corresponds to a probability of being correct of ≥75 per cent. Our results do not change appreciably if we instead adopt only those objects in VVDS with highly secure (>95 per cent) spectral redshifts. Galaxies suspected of containing an AGN due to either detection in X-rays, radio, broadband spectral features, or power-law infrared spectral energy distributions (SEDs) were kept in our final sample as the redshift of such sources still remains useful for a variety of purposes. Further, the presence of the AGN does not necessarily preclude the possibility that some or all of the physical parameters derived from our SED fitting process (see Section 2.4) remain valid as, in many cases, the rest-frame ultraviolet/optical/NIR portion of the SED remains dominated by the stellar component of such galaxies. Regardless, our results do not change appreciably if these objects are excluded from all analysis. The number of spectral targets, secure spectral redshifts, and with the total number of secure spectral extragalactic redshifts in the redshift range studied in this paper are given in Table 1. These numbers include both serendipitous detections and all redshifts from non-ORELSE surveys with the exception of VVDS. Also in Table 1 we list the properties of the known ORELSE groups and clusters listed in Hung et al. (2019). For a full account of the groups and clusters detected in the ORELSE fields, see Hung et al. (2019).

X = 0 is reserved for target galaxies, X = 1 for broadline AGNs, and X = 2 for non-targeted objects that fell serendipitously on a slit at a spatial location separable from the target. 2.3 CI1137+3000

The only ORELSE field included in the sample presented in this paper that has not appeared in detail in previous ORELSE studies is CI1137. The remainder of the fields as well as the associated observations on each field are discussed extensively in other ORELSE papers (e.g. Lemaux et al. 2012, 2017; Rumbaugh et al. 2017, 2018; Tomczak et al. 2017, 2019). The CI1137 field is unique among the ORELSE sample as it is the only cluster selected at radio wavelengths. This selection was based on the presence of a large (~240 kpc) wide-angle tailed radio source (WAT; Owen & Rudnick 1997) detected in the Very Large Array (VLA) Faint Images of the Radio Sky at Twenty Centimeters (Becker, White & Helfand 1995) survey, which is indicative of a Fanaroff–Riley type I (Fanaroff & Riley 1974) active galactic nucleus (AGN) or a similar phenomenon interacting at high relative velocities with ambient material, typically that formed by the presence of a group or cluster (O’Donoghue, Eilek & Owen 1993; Blanton et al. 2003). An optical/NIR imaging and Keck I/Low Resolution Imaging Spectrometer (LRIS; Oke et al. 1995) spectroscopic campaign in the region surrounding the location of the WAT confirmed the presence of a structure at a systemic redshift of (z) = 0.9559 with an LOS galaxy velocity dispersion of σ_{gal} = 530^{100}_{90} \text{ km s}^{-1} based on 10 members (Blanton et al. 2003). The projected distance of the WAT with respect to the structure centre as determined by the ORELSE data (see below), ~0.25 h_{70}^{-1}, is well within R_{500} and consistent with the typical observed location of WATs in lower redshift clusters making it extremely likely that its origins result from interaction with a dense medium.

Both proprietary and archival optical/NIR imaging were collected on the CI1137 field from Subaru/Suprime-Cam, UKIRT/WFCAM, and Spitzer/IRAC to the depths listed in Table 2. These data were
The parameter "\(z_{\text{peak}}\)" was adopted as the measure of \(z_{\text{phot}}\) with the uncertainties on this parameter estimated from the PDF of each source. Additional fitting for each source was done to the Pickles (1998) stellar library and was used, in conjunction with a variety of other criteria, to create a ‘use flag’ for each source (see Tomczak et al. 2017 for more details on this flag). Objects with a use flag set to zero, except for those objects which were spectroscopically determined to be a star at high confidence, were removed from all subsequent analysis. These objects typically totalled \(\sim 5\)–10 per cent of the total number of objects in a given ORELSE field. The precision and accuracy of the photometric redshifts were estimated from fitting a Gaussian to the distribution of \((z_{\text{spec}} - z_{\text{phot}})/(1 + z_{\text{spec}})\) measurements in the range \(0.5 \leq z \leq 1.2\) and a use flag = 1. This exercise resulted in a precision of \(\sigma_{\Delta z}/(1 + z) = 0.017 - 0.038\), with an average of 0.029, and a catastrophic outlier rate \((\Delta z/(1 + z) > 0.15)\) that ranges from \(\eta \sim 1.7\)–9.5 per cent, with an average of 5.8 per cent, across all ORELSE fields to limit of \(i \leq 24.5\). At this stage, if applicable, systematic offsets from zero in the \((z_{\text{spec}} - z_{\text{phot}})/(1 + z_{\text{spec}})\) distribution in each field were catalogued and corrected by applying the opposite of this offset to all \(z_{\text{phot}}\) values. These offsets ranged from \(-0.006\) to 0.005 with a median offset of 0.001. Subsequently, an additional phase of SED fitting was run with EAZY in which either the high-Q \(z_{\text{spec}}\) when available, or the \(z_{\text{phot}}\) was set as a redshift prior in order to estimate extinction-uncorrected rest-frame magnitudes estimated from the best-fitting template following the methodology of Brammer et al. (2011).

These rest-frame magnitudes are primarily used in this paper to determine, in a binary fashion, the state of star formation in objects that enter our final sample. For all such objects, we adopt the rest-frame \(M_{\text{NUV}} - M_{\text{K}}\) versus \(M_{\text{K}} - M_{\text{J}}\) cuts used in Lemaux et al. (2014b) to delineate between galaxies whose rest-frame colours are dominated by older and younger stellar populations used in Lemaux et al. (2014b). These two colours, or similar proxies, have been found to be intimately tied to the level of star formation activity (Arnouts et al. 2007; Martín et al. 2007; Wyder et al. 2007) and dust content of galaxies (e.g. Arnouts et al. 2013) and are extremely effective in separating the two populations independent of dust concerns (e.g. Williams et al. 2009; Ilbert et al. 2010, 2013; Muzzin et al. 2013; Moutard et al. 2016b). We use the following redshift-dependent criteria to define objects dominated by older stellar populations:

\[
M_{\text{NUV}} - M_{\text{K}} \geq 2.8(M_{\text{K}} - M_{\text{J}}) + 1.51 \land \\
M_{\text{NUV}} - M_{\text{K}} \geq 3.75 \land (z < 1) \\
M_{\text{NUV}} - M_{\text{K}} \geq 2.8(M_{\text{K}} - M_{\text{J}}) + 1.36 \land \\
M_{\text{NUV}} - M_{\text{K}} \geq 3.60 \land (z \geq 1).
\]

Note that at all redshifts investigated in this study in all ORELSE fields there existed imaging in bands whose filter curves in the rest-frame appreciably overlapped with each of the three rest-frame bands used here such that template-based \(k\)-corrections are minimized. Galaxies meeting their redshift-appropriate criteria above were designated as 'quiescent', while all remaining galaxies were classified as ‘star-forming’.

As a rough estimate of the level of \(S_{\text{SFR}}\), a classification of quiescent/star-forming corresponds to, in Ilbert et al. (2013) it is shown that the quiescent region of this diagram, defined in a similar way to our own, contains nearly all galaxies with \(\log(S_{\text{SFR}}[S_{\text{ED}}] < 10^{-11}\) for the redshift range considered in this study. While we were not able to estimate \(S_{\text{SFR}}\) values from our DEIMOS/LRIS spectroscopy for galaxies classified as quiescent due to fear that the main emission line present in our spectroscopy across the full

### Table 2. Photometry.

<table>
<thead>
<tr>
<th>Filter</th>
<th>Telescope</th>
<th>Instrument</th>
<th>Depth$^a$</th>
</tr>
</thead>
<tbody>
<tr>
<td>CI1137</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>B</td>
<td>Subaru</td>
<td>Suprime-Cam</td>
<td>25.9</td>
</tr>
<tr>
<td>V</td>
<td>Subaru</td>
<td>Suprime-Cam</td>
<td>26.0</td>
</tr>
<tr>
<td>$R_C$</td>
<td>Subaru</td>
<td>Suprime-Cam</td>
<td>25.3</td>
</tr>
<tr>
<td>$I_C$</td>
<td>Subaru</td>
<td>Suprime-Cam</td>
<td>24.4</td>
</tr>
<tr>
<td>$I^+$</td>
<td>Subaru</td>
<td>Suprime-Cam</td>
<td>25.1</td>
</tr>
<tr>
<td>$Z^+$</td>
<td>Subaru</td>
<td>Suprime-Cam</td>
<td>24.6</td>
</tr>
<tr>
<td>$J$</td>
<td>UKIRT</td>
<td>WFCAM</td>
<td>22.5</td>
</tr>
<tr>
<td>$K$</td>
<td>UKIRT</td>
<td>WFCAM</td>
<td>21.6</td>
</tr>
<tr>
<td>[3.6]</td>
<td>Spitzer</td>
<td>IRAC</td>
<td>23.2</td>
</tr>
<tr>
<td>[4.5]</td>
<td>Spitzer</td>
<td>IRAC</td>
<td>23.3</td>
</tr>
</tbody>
</table>

$^a$80 per cent completeness limits derived from the recovery rate of artificial sources inserted at empty sky regions.

reduced, and PSF-matched photometry was measured for all 10 bands in a manner identical to that presented in Section 2.1 and Tomczak et al. (2017) using the $Z^+$-band image as the detection image. In total, six masks were observed with the Keck ii/DEIMOS from 2011 January to 2015 April as part of the ORELSE survey. Integration times varied from 2600 to 14 400 s per mask$^{10}$ under mostly photometric conditions and seeing which varied from 0.5 to 1.2 arcsec. All observations were taken with the 1200 l mm$^{-1}$ grating tilted to a central wavelength of 7850 Å and used the OG550 order blocking filter. Spectroscopic targets were assigned following the scheme presented in Lubin et al. (2009), with objects that had $R_CI^+$ $Z^+$ colours consistent with those expected for a generic cluster red sequence at $z \sim 0.95$ were highly prioritized (i.e. priority 1). In the CI1137 field, out of 827 targets, only 116 (14.0 per cent) had colours consistent with our priority 1 colour window (i.e. $1.0 \leq R_C - I^+ \leq 1.4, 0.6 \leq I^+ - Z^+ \leq 1.0$). In total, these observations yielded 539 (438) high-quality (extragalactic) redshifts, of which 66 fell in the redshift range of the main LSS of $0.937 < z < 0.963$. Only one massive structure was found in this field, a structure centred at $[\alpha_{2000}, \delta_{2000}] = [174.397 86, 30.008 93]^{11}$ with a systemic redshift of $z = 0.9553$ and a LOS galaxy velocity dispersion of $\sigma_v = 535 \pm 81$ km s$^{-1}$ based on 28 members within 1 h$^{-1}$ Mpc from the structure centre. Following the formalism of Lemaux et al. (2012), these measurements indicate a structure with a virial mass of $\log(M_{\text{vir}}/M_{\odot}) = 14.1 \pm 0.2$, consistent with the presence of either a low-mass cluster or a high-mass group.

### 2.4 Spectral energy distribution fitting and stellar mass limits

Fitting of the SED of each object detected in the detection image of each field was first performed on aperture magnitudes measured on the PSF-homogenized images using the Easy and Accurate $z_{\text{phot}}$ from Yale (EAZY; Brammer, van Dokkum & Coppi 2008) code. This fitting was done for the purposes of estimating photometric redshifts (hereafter $z_{\text{phot}}$) of each object and an associated probability distribution function (PDF). The PDFs were generated by minimizing the $\chi^2$ of the observed flux densities and a set of basis templates at each redshift (including emission lines, see Brammer et al. 2011). See Tomczak et al. (2017) for more details on this fitting process.

$^{10}$The shallowest mask was used primarily to select targets for subsequent masks. The next shortest integration time was 7200 s.

$^{11}$This centre represents the $Z^+$ luminosity-weighted centre of all spectroscopic member galaxies. For more details on this calculation, see Ascaso et al. (2014).
The procedure to estimate SSFR of the galaxies in the spectral sample classified as (see below), and the Wuyts et al. (2013) relation was adopted to the brightness of an exponentially declining BC03 model galaxy at Tomczak et al. (2017). Stellar mass limits are determined by scaling for all ORELSE fields using a method similar to that described in over the redshift range studied here. Here we test that statement \( \leq 0.55 \) to an SED with photometry that has been tweaked by a Gaussian and uncertainties are derived through 100 realizations of re-fitting. Exponentially declining stellar population synthesis Bruzual & Charlot (2003) models (hereafter BC03) were adopted with a Chabrier (2003) initial mass function, a Calzetti et al. (2000) extinction law, and a stellar-phase metallicity fixed at \( Z = Z_\odot \). The ranges of allowed parameters are identical to those of Tomczak et al. (2017), with the imposition that the age of a given objects could not exceed the age of the universe at the estimated redshift of that object. The value of each parameter is taken from the best-fit value and uncertainties are derived through 100 realizations of re-fitting to an SED with photometry that has been tweaked by a Gaussian random multiple of its photometric errors for each band (as in, e.g. Ryan et al. 2014).

2.5 Spectral representativeness

As discussed in Tomczak et al. (2017), the ORELSE imaging of the fields presented in that paper is broadly to a sufficient depth to detect all galaxy types to a stellar mass limit of \( \log(\mathcal{M}_\star/\mathcal{M}_\odot) > 10 \) over the redshift range studied here. Here we test that statement for all ORELSE fields using a method similar to that described in Tomczak et al. (2017). Stellar mass limits are determined by scaling the brightness of an exponentially declining BC03 model galaxy at each redshift, in steps of \( \Delta z = 0.1 \), to the 80 per cent magnitude completeness limits of the detection image in each field. The BC03 template was generated with a formation redshift of \( z_f = 5 \), a star-formation history (SFH) with an e-folding time (\( \tau \)) of 1 Gyr, a Chabrier (2003) IMF, and no internal dust extinction. The scaling factor used to match the brightness of the model galaxy at each redshift to the magnitude completeness limits is similarly used to scale the stellar mass of the model galaxy that then sets the mass completeness limit. Note that this template will contain a higher stellar mass-to-light ratio than the vast majority of galaxies at all redshifts studied here and, thus, represents a conservative upper limit to the stellar mass completeness of the ORELSE images. We find that the statement made in Tomczak et al. (2017) generally holds for the full ORELSE sample presented here, with most fields having stellar mass limits of \( \log(\mathcal{M}_\star/\mathcal{M}_\odot) \sim 10 \) or below from 0.55 \( \leq z \leq 1.4 \). While a few fields have stellar mass limits in excess of \( 10^{10}\mathcal{M}_\odot \) at or near the highest redshifts presented in this study, our results do not change meaningfully if we tailor the redshift range included in each field to set comparable stellar mass limits across the entire sample. As such, we ignore any effects of the differential loss to our sample of lower mass, older quiescent galaxies in a few of the ORELSE fields at the highest redshifts studied in this paper. Note that extremely dust-reddened galaxies at these stellar masses broadly comparable rest-frame optical/UV fluxes to galaxies dominated by a near-maximal age stellar population. As the detection bands\(^{13}\) employed in ORELSE probe the rest-frame optical or redder-portion of the UV for galaxies in the redshift range considered in this study, such dusty populations should generally be detected by our imaging. Such a statement is corroborated by the routine detection of dusty star-forming galaxies in both our imaging and spectroscopy (see e.g. Kocevski et al. 2011; Shen et al. 2017).

In Shen et al. (2017), a spectral sample drawn from five ORELSE LSSs was compared to an underlying photometric population at various stellar mass and colour limits. This comparison was made in order to understand to what limits the spectral sample in that study could be considered representative of the galaxy population at these redshifts as a whole. It was found for an apparent magnitude cut of \( i \leq 24.5 \) that the ORELSE spectral sample presented in that paper could be broadly characterized as being representative of the underlying photometric population in the range \( \log(\mathcal{M}_\star/\mathcal{M}_\odot) \geq 10.2 \) and \( M_{\text{NUV}} - M_i \geq 2.5 \). Here we expand this comparison to all of ORELSE following a method similar to that presented in Shen et al. (2017). Spectral and photometric samples from all ORELSE fields were selected within the redshift range \( 0.55 \leq z_{\text{spec/phot}} \leq 1.4 \) and to an apparent magnitude limit of \( 18.5 \leq i \leq 24.5 \) or \( 18.5 \leq z \leq 24.5 \) depending on whether the object is at \( z \leq 1 \) or \( z \geq 1 \), respectively. The \( z_{\text{phot}} \) sample in each ORELSE field consisted only of those objects without a high-\( Q \) \( z_{\text{spec}} \) that were within 60 arcsec of an object targeted for spectroscopy, a scheme which roughly selects all those objects which were potential spectroscopic targets which were not targeted or which were targeted but for which we had no secure \( z_{\text{spec}} \). The value 60 arcsec was chosen to ensure that the \( z_{\text{phot}} \) objects selected broadly shared the same depth and breadth of imaging as the actual spectral targets, but large enough that it does not exclude objects that fell within the spectral footprint that were separated from the spectral targets only by virtue of the chosen slit geometry (i.e. 60 arcsec is roughly half of the maximum interslit distance for our masks). In addition, only those objects with estimated stellar masses and rest-frame magnitudes and with a use flag = 1 (see Section 2.4) were included in this comparison. Rest-frame \( NUVrJ \) colour–colour diagrams for these spectroscopic and \( z_{\text{phot}} \) samples are shown in Fig. 2 along with the delineation lines for quiescent and star-forming galaxies described earlier in this section. Note that the upper right portion of the star-forming locus, i.e. objects with \( M_{\text{NUV}} - M_i > 1 \) and \( M_{\text{NUV}} - M_i > 3.5 \), a region where the dustiest galaxies are thought to lie (e.g. Arnouts et al. 2013) are well represented in both samples. Further, the relative abundance of objects in this region of colour–colour space is comparable to legacy fields like, e.g. COSMOS, where the detection band(s) employed are well in the rest-frame NIR at these redshifts where dust effects are minimal (e.g. Ilbert et al. 2013; Laigle et al. 2016).

We compared the stellar mass and rest-frame \( M_{\text{NUV}} - M_i \) colour distributions for the spectroscopic and \( z_{\text{phot}} \) samples in bins of 0.2 dex and 0.5 mag, respectively, by means of a Kolmogorov–Smirnov

\(^{13}\)See Tomczak et al. (2017), Rumbaugh et al. (2018), and Section 2.3 for details on detection bands for each ORELSE field.
Through these comparisons, in conjunction with the knowledge that the imaging in all fields is of sufficient depth to detect galaxies of all types to \( \log(M_*/M_\odot) \geq 10 \), we conclude that the ORELSE spectral sample shown here is broadly representative of the full galaxy population in the range \( \log(M_*/M_\odot) \geq 10.0, M_{\text{NUV}} - M_r > 2.0 \), and \( 0.55 \leq z \leq 1.4 \). These limits, along with the associated apparent magnitude and quality flag cuts, set the final spectral sample that is presented in the remainder of the paper. This final sample consists of 4552 galaxies with secure spectral redshifts and reliable photometry. Since we have shown that this sample is representative of the full galaxy population in the range of colours, stellar masses, and redshifts we are probing, and, further, since \( z_{\text{phot}} \) objects have relatively large uncertainties in all three of the parameters were of interest in \( [z, M_*, \log (1 + \delta_{\text{gal}})] \), we chose for the remainder of the paper to exclusively use this final spectral sample. We note that none of the results presented in this paper change meaningfully if we slightly change the stellar mass and colour limits that define our final sample (i.e. by \( \pm 0.2 \) dex and \( \pm 0.5 \) mag, respectively).

### 2.6 Local overdensity

In order to estimate the local environment of the galaxies in our sample, we employ the Voronoi Monte Carlo (VMC) technique described in Lemaux et al. (2017). This estimator has been employed in a variety of different works at both low and high redshift (e.g. Shen et al. 2017, 2019; Tomczak et al. 2017; Pelliccia et al. 2018; Cucciati et al. 2018; Lemaux et al. 2018, 2017; Tomczak et al. 2019), and its precision and accuracy are discussed briefly in Tomczak et al. (2017). Further quantitative measures of the precision and accuracy of this reconstruction for the detection and characterization

---

**Figure 2.** Left: Rest-frame \( M_{\text{NUV}} - M_r/M_* - M_J \) colour–colour diagram (CMD) for all objects in ORELSE in the photometric redshift range indicated that fall within \( \lesssim 60 \) arcsec of a spectral target and have an apparent magnitude in the range \( 18.5 \leq i \leq 24.5 \) and \( 18.5 \leq z \leq 24.5 \) for objects with \( z_{\text{phot}} < 1 \) and \( z_{\text{phot}} \geq 1 \), respectively. The number of \( z_{\text{phot}} \) objects satisfying these criteria is given in the top left. The colour bar in the bottom right of the panel indicates the associated colour for the number of objects in each two-dimensional bin. The solid line shows an average of the colour–colour criteria used to delineate quiescent galaxies (top left region, above the lines) and star-forming galaxies (bottom right region, below the lines) at \( z < 1 \) and \( z \geq 1 \) (see Section 2.4). Right: The same as the left-hand panel, but for galaxies spectroscopically confirmed in the redshift range indicated subject to the same apparent magnitude cuts as the \( z_{\text{phot}} \) objects.
Figure 3. Top Left: Rest-frame $M_{NUV} - M_r$ colour-stellar mass diagram (CSMD) for all objects in ORELSE in the photometric redshift range indicated in Section 2.5 that fall within $<60$ arcsec of a spectral target and have an apparent magnitude in the range $18.5 \leq i \leq 24.5$ and $18.5 \leq z \leq 24.5$ for objects with $z_{\text{phot}} < 1$ and $z_{\text{phot}} \geq 1$, respectively. The number of $z_{\text{phot}}$ objects satisfying these criteria is given in the top left. The colour bar in the bottom right of the panel indicates the number of objects in each two-dimensional bin. Top Right: The same as the left-hand panel, but for galaxies spectroscopically confirmed in the redshift range indicated subject to the same apparent magnitude cuts as the $z_{\text{phot}}$ objects in the left-hand panel. Bottom: Fraction of $z_{\text{spec}}$ to $z_{\text{phot}}$ objects from the top two panels as a function of rest-frame colour and stellar mass. At least 10 $z_{\text{phot}}$ objects are required to fall within a given colour/stellar mass bin in order to be plotted. Accompanying histograms of the ORELSE $z_{\text{spec}}$ and $z_{\text{phot}}$ samples in stellar mass and rest-frame colour are shown and are normalized such that the area of two histograms for each parameter are equal for stellar masses in excess of $10^{10} M_\odot$ and colours bluer than $M_{NUV} - M_r = 2$. As can be seen in both the one- and two-dimensional histograms, the fraction of spec-zs in ORELSE drops dramatically below these two limits. Above these limits, the spectral sample is found to be broadly representative of the underlying $z_{\text{phot}}$ population in terms of these two parameters (see Section 2.5).

of groups and clusters will be presented in Hung et al. (2019). Here, the VMC method is described briefly.

The VMC attempts to combine the information provided by the high-precision spectroscopic redshifts with the full $P(z)$ information contained within each object without a spectroscopic redshift to determine the local overdensity in thin redshift slices that span the redshift range where both our spectroscopic and photometric data are sensitive ($0.55 \leq z \leq 1.4$). In this scheme, all objects with secure spectroscopic redshifts are considered to be at the measured redshift. For those objects without a secure spectroscopic redshift, multiple realizations are run for each redshift slice to sample from the statistics of the photometric redshift assigned to each object to determine the range of possible density fields. In practice, for each Monte Carlo realization, Gaussian sampling is performed to determine a new set of $z_{\text{phot}}$ values for each object without a high quality $z_{\text{spec}}$ (but with a good use flag, see Section 2.4) for that realization. The sampled value, in units of $\sigma$, is then multiplied by either the effective $1\sigma$ lower or upper uncertainty on $z_{\text{phot}}$ for that object depending on which side of the peak of the Gaussian sample fell. This value is then either subtracted from or added to the original $z_{\text{phot}}$ for each object to create a new set of $z_{\text{phot,MC}}$ for that realization. A thin redshift slice is cut from the combined
catalogue which includes both objects with new \( z_{\text{phot}} \) values as well as all galaxies with high quality extragalactic \( z_{\text{spec}} \), and Voronoi tessellation is performed for the realization of that slice. This process is performed, in total, for 100 realizations for each of the 85 redshift slices running from 0.55 \( \leq z \leq 1.4 \), which is the number of slices required to span this redshift range for a slice width of \( \pm 1500 \text{ km s}^{-1} \) and steps between the central redshift of adjacent slices that are half of the slice depth (i.e. 1500 km s\(^{-1}\)).

For each realization of each slice, a grid of \( 75 \times 75 \) kpc is created to sample the underlying local density distribution. The local density at each grid value for each realization and slice is set equal to the inverse of the Voronoi cell area (multiplied by \( D^3 \)) of the cell that encloses the central point of that grid. Final local densities, \( \Sigma_{\text{VMC}} \), for each grid point in each redshift slice are then computed by median combining the values of the 100 realizations of the Voronoi maps for that slice. The local overdensity value for each grid point is then computed as \( \log(1 + \delta_{\text{gal}}) = \log(1 + (\Sigma_{\text{VMC}} - \Sigma_{\text{VMC}})/\Sigma_{\text{VMC}}) \), where \( \Sigma_{\text{VMC}} \) is the median \( \Sigma_{\text{VMC}} \) for all grid points over which the map is defined (i.e. where there is coverage in a sufficient number of imaging bands). By adopting local overdensity rather than local density as a proxy of environment, we largely mitigate issues of sample selection and differential bias as a function of redshift. For VMC overdensity maps in all ORELSE fields, photometric and spectroscopic catalogues were cut at 18 \( \leq z \leq 24.5 \) or 18 \( \leq z \leq 24.5 \) depending on the field. Uncertainties associated with the overdensity estimate for each object are taken from the 16th and 84th percentile of the \( \log(1 + \delta_{\text{gal}}) \) distribution generated from all 100 Monte Carlo iterations.

\section{THE PERSISTENCE OF THE COLOUR–DENSITY RELATION AND EFFICIENT ENVIRONMENTAL QUENCHING TO \( z \sim 1.4 \)}

\subsection{The quiescent fraction in ORELSE}

In Fig. 5 we show the quiescent fractions, \( f_q \), of galaxies in our final spectral sample as a function of \( z \), \( \log(1 + \delta_{\text{gal}}) \), and \( \mathcal{M}_s \). The quantity \( f_q \) is simply defined as \( N_q/(N_S + N_q) \), where \( N_q \) and \( N_S \) refer to the number of quiescent and star-forming galaxies, respectively. Galaxies are separated into star-forming and quiescent types based on the scheme described in Section 2.4. As noted in Section 2.5, this sample is comprised only of those ORELSE galaxies with \( \log(M_s) \geq 10 \) and \( M_{\text{NUV}} - M_r \geq 2 \) and, as such, represent an upper limit to the true quiescent fractions of all galaxies with \( \log(M_s) \geq 10 \) in the redshift range 0.55 \( \leq z \leq 1.4 \). However, as shown in Fig. 3, galaxies with colours bluer than \( M_{\text{NUV}} - M_r = 2 \) only begin to comprise a non-negligible fraction (~15\% of the total galaxy population for stellar masses in the range 10.0 \( \leq \log(M_s) \leq 10.5 \)). For higher stellar masses, the contribution of galaxies at such colours is effectively negligible (~1\%). We choose not to make a correction for this effect throughout the paper as we do not know a priori the relative number of galaxies with \( M_{\text{NUV}} - M_r < 2 \) in each redshift and environmental bin.

The full sample is broken into three different \( \mathcal{M}_s \) bins for each redshift and environmental density bin that broadly follow the ranges 10–10.5 (low-\( \mathcal{M}_s \)), 10.5–10.9 (med-\( \mathcal{M}_s \)), and >10.9 (high-\( \mathcal{M}_s \)) in \( \log(M_s/M_\odot) \). Similarly, galaxies are broken into three different \( \log(1 + \delta_{\text{gal}}) \) bins for each that broadly follow the ranges <0.3 [low-log (1 + \( \delta_{\text{gal}} \)), 0.3–0.7 [intermediate-log (1 + \( \delta_{\text{gal}} \)], and >0.7 [high-log (1 + \( \delta_{\text{gal}} \)]. Physically, these three environmental density bins can be thought of as corresponding to field galaxies, the dense field combined with infall or intermediate regions of groups and the infall regions of clusters, and the intermediate to core regions of both groups and clusters, respectively. These interpretations are based on the distribution of galaxies in each local overdensity bin in projected radial and differential velocity space relative to the closest group or cluster (i.e. global density \( \equiv R_{\text{proj}}/R_{200} \times \delta_v/\sigma_v \)). For precise definitions of infall, intermediate, and core regions as well as our method for calculating global density see Pelliccia et al. (2018), Shen et al. (2019), and references therein. It is important to note also that our VMC log (1 + \( \delta_{\text{gal}} \)) metric also shows a significant (~3\%) degree of correlation with the global metric referred to above. This degree is also for the more invertebrate environmental metric that simply employs the normalized projected radial distances of galaxies relative to the centroid of their nearest group or cluster. In terms of this metric, the high-log (1 + \( \delta_{\text{gal}} \)) bin galaxies fall within 0 \( \leq R_{\text{proj}}/R_{200} \leq 1 \), with an average \( R_{\text{proj}}/R_{200} = 0.58 \). Since all ORELSE groups and clusters are included in this analysis, the average total mass of structures considered here is \( \log(M_{\text{clus}}/M_\odot) = 14.4 \). The exact limits for both the stellar mass and environmental density bins at each redshift are slightly modulated (±0.1 dex) for each bin in order to attempt to roughly impose equal numbers of galaxies in each \( \tau \log(1 + \delta_{\text{gal}})/M_s \), bin. Uncertainties associated with each quiescent fraction are calculated using the formula \( N_q/N_S(N_S + N_q)^2 \) derived from Poisson statistics.

The general relationship of \( f_q \) with respect to each of the three parameters against it is plotted, which is apparent from a visual inspection of Fig. 5. In essentially all epochs observed in ORELSE over the redshift range studied here (0.55 \( \leq z \leq 1.4 \)), galaxy populations with increasing stellar masses are observed with higher quiescent fractions. The behaviour of \( f_q \) with redshift across all environmental density bins appears to be broadly self-similar, with med- to high-\( M_s \) galaxies showing little to no significant decrease in their \( f_q \) over the observed redshift range and low-\( M_s \) galaxies exhibiting a clear monotonic decrease in \( f_q \) with redshift.

The behaviour of \( f_q \) appears similar for galaxies residing in each type of environment studied here, there is a clear increase in \( f_q \) of 0.1–0.2 at essentially all stellar masses and redshifts when moving from the field to galaxies primarily inhabiting group/cluster environments [i.e. low- to high-log (1 + \( \delta_{\text{gal}} \)]. It appears from these measurements that the ability of high-density environments to quench galaxies of essentially all stellar masses persists from 0.55 \( \leq z \leq 1.4 \), a claim that is consistent with other works studying the galaxy populations of groups and clusters at similar redshifts (e.g. Cooper et al. 2007, 2010; Cacciari et al. 2010, 2017; Kovač et al. 2010; Vulcani et al. 2010; Muzzin et al. 2012; Balogh et al. 2016; Nantais et al. 2017). A possible exception to this statement enters for the low-\( M_s \) sample at the highest redshifts observed in ORELSE. At these redshifts and stellar masses the \( f_q \) of galaxies in the high-log (1 + \( \delta_{\text{gal}} \) bin becomes formally consistent with that of their counterparts inhabiting field environments. This lack of a significant difference implies that group/cluster environments at \( z \geq 1 \) are either marginally effective or completely ineffectual at quenching galaxies at stellar masses of \( \log(M_s/M_\odot) \sim 10.2 \).

In Table 3 we list \( f_q \) values and their associated uncertainties for all environmental, redshift, and stellar mass bins measured in ORELSE. Using the values in Table 3 we performed a \( \chi^2 \) fit of \( f_q \) as a function of redshift, local environment, and stellar mass following.
the form:

\[ f_q(z, M_\star, \log(1 + \delta_{\text{gal}})) = \alpha z^\xi + \beta \log(M_\star/M_\odot)^\gamma + \gamma \log(1 + \delta_{\text{gal}})^\kappa, \]

which yielded best-fitting parameters of \( \alpha = -1.8544 \pm 0.0002, \beta = 8.3876 \pm 0.0001, \gamma = -11.8321 \pm 0.0002, \xi = 0.1421 \pm 0.0007, \) and \( \gamma = 0.2224 \pm 0.0007, \) and \( \kappa = -0.0076 \pm 0.0005, \) where the uncertainties are taken from the covariance matrix of the fit. This formula is not specific to ORELSE and is generally applicable to the galaxy populations in the range of redshifts \( (0.55 \leq z \leq 1.4) \), stellar masses \( (10 \leq \log(M_\star/M_\odot) \leq 11.6) \), and overdensity values \( (0.1 \leq \log(1 + \delta_{\text{gal}}) \leq 2.0) \) studied here. We stress that this formula should not be applied outside these ranges. Further, we remind the reader that this formula will tend to slightly overestimate \( f_q \) by perhaps as much as \( 0.1 \) for galaxies in the lower stellar mass range in our sample \( [i.e. \log(M_\star/M_\odot) < 10.5] \) due to the way in which our sample was created (see Section 2.5).

3.2 Environmental quenching efficiency

Previous studies of the relationship of \( f_q \) with various galaxy properties, including the environment in which galaxies reside, have attempted to quantify the efficacy of environmentally related processes in quenching the star formation of member galaxies by defining a parameter known as the ‘conversion factor’ or ‘environmental quenching efficiency’. This parameter, which we denote \( \Psi_{\text{convert}} \), is defined as the number of galaxies, usually at a fixed stellar mass, which have converted from star-forming to quiescent as a result of entering and interacting within a group or cluster environment. Practically, \( \Psi_{\text{convert}} \) is usually defined as the difference in the \( f_q \) fractions between the highest and lowest density bins in a given sample divided by the fraction of star-forming galaxies in the lower density sample, i.e. \( \Psi_{\text{convert}} = (f_q, \text{high} - \log(1 + \delta_{\text{gal}}) - f_q, \text{low} - \log(1 + \delta_{\text{gal}}))/f_{\text{SF}, \text{high}} - f_{\text{SF}, \text{low}}(1 + \delta_{\text{gal}}) \) (van den Bosch et al. 2008; McGee, Bower & Balogh 2014). If the highest density bin in a given sample primarily contains group/cluster galaxies and the galaxies in the lowest density bin can be used as a proxy for a progenitor sample for those galaxies in the highest density bin, then \( \Psi_{\text{convert}} \) yields the fraction of star-forming galaxies which have converted to quiescent during coalescence into the group/cluster. These values can then be compared to accretion histories of groups/clusters as measured from \( N \)-body simulations (as in, e.g. McGee et al. 2009; Balogh et al. 2016) to determine the time-scale required to convert the galaxy from star-forming to quiescent, \( t_{\text{convert}} \), from the time it leaves the field for the group/cluster environment until the time its rest-frame colours place it in the quenched region of the \( NUVrJ \) diagram.

However, as discussed in Appendix A, such a concept when applied directly to the data is problematic due to a variety of issues such as the evolution of the quiescent fraction in the field, differential stellar mass growth, and differential sampling issues. To circumvent these issues we have created simulated galaxy populations with which to compare the \( f_q \) measurements of the group/cluster galaxies observed in the ORELSE data. These simulations, meant to simulate the \( f_q \) of group and cluster galaxies in the absence of environmental quenching processes, are described in detail in Appendix A. The simulations are performed 200 times, generating 200 unique sets of simulated group/cluster galaxies in the redshift and stellar mass ranges \( 0.65 \leq z \leq 4 \) and \( \log(M_\star/M_\odot) \geq 9 \), respectively. The accretion rate of the cluster in each simulation at each redshift is based on the halo accretion history presented in McGee et al. (2009) for a \( M_{\text{halo}, z=0}/M_\odot \geq 14.5 \) cluster. This accreted mass is transformed into a galaxy population through the halo-to-stellar mass relation presented in Moster, Naab & White (2013) and galaxies are probabilistically assigned a state of quiescence or star-forming based on statistics of the galaxy stellar mass function (SMF) of Tomczak et al. (2014). For those galaxies assigned a star-forming state, an exponentially declining SFR is ascribed.
Persistence of environmental quenching to $z \sim 1.4$

Figure 5. **Top left:** Quiescent fraction, $f_q$, of galaxies in the lowest overdensity regime of the final spectral sample [$\log(1+\delta_{\text{gal}}) \lesssim 0.3$, i.e. field environments] as a function of redshift in three different stellar mass ranges (roughly $10^{10.45}$, $10.45-10.9$, and $>10.9$ in $\log(M_*/M_\odot)$ for the low-, med-, and high-$M_*$ samples, respectively). The median stellar mass of the galaxies in each stellar mass bin across all redshifts is shown in the top right. The apparent magnitude, stellar mass, and rest-frame colour criteria imposed on the final spectral sample is shown in the bottom left, and the number of galaxies in this environmental bin is given in the top left. Uncertainties are derived from Poissonian statistics. **Top right:** Identical to the top left panel but for the galaxies in the intermediate overdensity sample (i.e. infall/intermediate regions of groups/clusters). **Bottom:** Identical to the other two panels but for galaxies in the highest densities in ORELSE (i.e. intermediate/core regions of groups/clusters). Galaxy populations in higher density environments and at higher stellar masses are generally seen to exhibit a higher incidence of quiescent galaxies from $0.55 \leq z \leq 1.4$. Only the low-$M_*$ galaxies appear to evolve strongly in $f_q$ with redshift, with that of galaxies in the other two stellar mass bins showing moderate to no redshift evolution.

to each galaxy with $\tau = 2.0$ Gyr and a normalization set by sampling from the relationship between $SFR-M_*$ for star-forming galaxies presented in Tomczak et al. (2016). The average fraction of quiescent galaxies and the normalized median absolute deviation (NMAD; Hoaglin, Mosteller & Tukey 1983) of the $f_q$ distributions in the same three stellar mass bins applied to our final spectral sample (e.g. low-, med-, and high-$M_*$, see Section 3.1) set the values for the low-log ($1+\delta_{\text{gal}}$) populations in the calculation of $\Psi_{\text{convert}}$. As a reminder, these $f_q$ values are measured on simulated cluster galaxy populations in each stellar mass and redshift bin that have been unaffected by environmental quenching for their entire evolution beyond what is experienced in field samples. The behaviour of $f_q$ for simulated galaxies with stellar mass and redshift is shown in Fig. 6.

The simulated $f_q$ values are then used in conjunction with the observed $f_q$ values measured for the different galaxy populations in the high-log ($1+\delta_{\text{gal}}$) bin to estimate the environmental quenching efficiency in the ORELSE survey as a function of redshift and stellar mass. In other words, the simulated $f_q$ values are used in this calculation instead of our observed low-log ($1+\delta_{\text{gal}}$) $f_q$ values as the former are almost certainly more appropriate when calculating quenching efficiencies (see Appendix A). The errors on the quenching efficiency are calculated from a combination of the errors on the $f_q$ values measured for high-log ($1+\delta_{\text{gal}}$) galaxies and the error on the median $f_q$ of all simulations. These environmental quenching efficiencies, plotted in Fig. 7, exhibit a behaviour that largely mirrors that observed in the $f_q$ values shown in Fig. 5. Two specific parallels can be made. First, the quenching efficiency decreases precipitously with increasing redshift for low-$M_*$ galaxies and becomes formally consistent with group/cluster environments having no transformative power at the highest redshifts observed in ORELSE ($z \sim 1.4$). Conversely, for galaxies above $\log(M_*/M_\odot) \geq 10.4$, $\Psi_{\text{convert}}$ appears to decrease gradually
or remain constant with increasing redshift. In other words, it appears that over the redshift range $0.55 \leq z \leq 1.4$, galaxy clusters and groups remain capable of efficiently quenching star formation activity in accreted field galaxies for $\log(\delta_{\text{gal}}) \geq 10.4$, but that groups and clusters at the higher end of the redshifts studied here ($z \geq 1.1$) are incapable of exerting themselves in the same way on galaxies with lower stellar masses.

The values of $\Psi_{\text{conv}}$, derived from galaxy populations in other surveys at a variety of different stellar masses and redshifts are overplotted in Fig. 7. The size of the points are set to small, medium, and large for galaxy samples with comparable stellar masses to those in the low-, med-, and high-$M_*$ bins in the ORELSE data. Though values associated with other surveys are generally subject to relatively large uncertainties, the measurements of $\Psi_{\text{conv}}$ made of the ORELSE samples appear generally consistent with other measurements at common redshifts and stellar masses. Marginally significant tension appears at $z \sim 1$ for med-$M_*$ galaxies with respect to the study of Fossati et al. (2017) based on the 3D-HST survey (Brammer et al. 2012) and the study of Nantais et al. (2017) based on galaxies selected from the Spitzer Adapation of the Red Sequence Cluster Survey (SpARCS; Muzzin et al. 2009; Wilson et al. 2009). Additional moderate tension appears for measurements of low-$M_*$ galaxies with respect to that measured from the Gemini Cluster Astrophysics Spectroscopic Survey (GCLASS; Muzzin

### Table 3. Quiescent fractions of the ORELSE spectral sample.

<table>
<thead>
<tr>
<th>Sample</th>
<th>$\bar{z}$</th>
<th>$f_q$</th>
<th>log($1+\delta_{\text{gal}}$)</th>
<th>log($\delta_{\text{gal}}/M_\odot$)</th>
<th>$N_{\text{gal}}$</th>
</tr>
</thead>
<tbody>
<tr>
<td>Low-$M_*/\text{low-log}(1+\delta_{\text{gal}})$</td>
<td>0.682</td>
<td>0.390 ± 0.035</td>
<td>0.162</td>
<td>10.24</td>
<td>195</td>
</tr>
<tr>
<td></td>
<td>0.789</td>
<td>0.333 ± 0.040</td>
<td>0.165</td>
<td>10.22</td>
<td>138</td>
</tr>
<tr>
<td></td>
<td>0.890</td>
<td>0.217 ± 0.030</td>
<td>0.121</td>
<td>10.26</td>
<td>189</td>
</tr>
<tr>
<td></td>
<td>1.041</td>
<td>0.156 ± 0.028</td>
<td>0.146</td>
<td>10.24</td>
<td>173</td>
</tr>
<tr>
<td></td>
<td>1.284</td>
<td>0.095 ± 0.045</td>
<td>0.088</td>
<td>10.27</td>
<td>42</td>
</tr>
<tr>
<td></td>
<td>0.686</td>
<td>0.475 ± 0.045</td>
<td>0.170</td>
<td>10.65</td>
<td>122</td>
</tr>
<tr>
<td></td>
<td>0.794</td>
<td>0.525 ± 0.046</td>
<td>0.172</td>
<td>10.61</td>
<td>120</td>
</tr>
<tr>
<td>Med-$M_*/\text{low-log}(1+\delta_{\text{gal}})$</td>
<td>0.901</td>
<td>0.394 ± 0.043</td>
<td>0.162</td>
<td>10.70</td>
<td>208</td>
</tr>
<tr>
<td></td>
<td>1.035</td>
<td>0.466 ± 0.036</td>
<td>0.131</td>
<td>10.65</td>
<td>189</td>
</tr>
<tr>
<td></td>
<td>1.262</td>
<td>0.361 ± 0.057</td>
<td>0.170</td>
<td>10.67</td>
<td>72</td>
</tr>
<tr>
<td></td>
<td>0.698</td>
<td>0.529 ± 0.061</td>
<td>0.192</td>
<td>11.10</td>
<td>68</td>
</tr>
<tr>
<td></td>
<td>0.791</td>
<td>0.402 ± 0.054</td>
<td>0.197</td>
<td>11.06</td>
<td>82</td>
</tr>
<tr>
<td>Med-$M_*/\text{med-log}(1+\delta_{\text{gal}})$</td>
<td>0.880</td>
<td>0.404 ± 0.065</td>
<td>0.192</td>
<td>11.13</td>
<td>57</td>
</tr>
<tr>
<td></td>
<td>1.064</td>
<td>0.630 ± 0.054</td>
<td>0.140</td>
<td>11.05</td>
<td>81</td>
</tr>
<tr>
<td></td>
<td>1.254</td>
<td>0.429 ± 0.084</td>
<td>0.123</td>
<td>11.06</td>
<td>35</td>
</tr>
<tr>
<td></td>
<td>0.692</td>
<td>0.441 ± 0.043</td>
<td>0.471</td>
<td>10.26</td>
<td>136</td>
</tr>
<tr>
<td></td>
<td>0.797</td>
<td>0.360 ± 0.046</td>
<td>0.441</td>
<td>10.26</td>
<td>111</td>
</tr>
<tr>
<td>Med-$M_*/\text{med-log}(1+\delta_{\text{gal}})$</td>
<td>0.896</td>
<td>0.343 ± 0.041</td>
<td>0.419</td>
<td>10.29</td>
<td>137</td>
</tr>
<tr>
<td></td>
<td>1.010</td>
<td>0.219 ± 0.042</td>
<td>0.418</td>
<td>10.24</td>
<td>96</td>
</tr>
<tr>
<td></td>
<td>1.260</td>
<td>0.045 ± 0.026</td>
<td>0.326</td>
<td>10.25</td>
<td>66</td>
</tr>
<tr>
<td></td>
<td>0.683</td>
<td>0.496 ± 0.042</td>
<td>0.463</td>
<td>10.69</td>
<td>141</td>
</tr>
<tr>
<td></td>
<td>0.798</td>
<td>0.532 ± 0.048</td>
<td>0.454</td>
<td>10.62</td>
<td>109</td>
</tr>
<tr>
<td>Med-$M_*/\text{high-log}(1+\delta_{\text{gal}})$</td>
<td>0.904</td>
<td>0.468 ± 0.052</td>
<td>0.426</td>
<td>10.72</td>
<td>171</td>
</tr>
<tr>
<td></td>
<td>1.032</td>
<td>0.419 ± 0.040</td>
<td>0.395</td>
<td>10.65</td>
<td>155</td>
</tr>
<tr>
<td></td>
<td>1.211</td>
<td>0.348 ± 0.059</td>
<td>0.375</td>
<td>10.70</td>
<td>66</td>
</tr>
<tr>
<td></td>
<td>0.695</td>
<td>0.609 ± 0.059</td>
<td>0.442</td>
<td>11.02</td>
<td>69</td>
</tr>
<tr>
<td></td>
<td>0.797</td>
<td>0.450 ± 0.050</td>
<td>0.453</td>
<td>11.03</td>
<td>100</td>
</tr>
<tr>
<td>Med-$M_*/\text{med-log}(1+\delta_{\text{gal}})$</td>
<td>0.898</td>
<td>0.475 ± 0.064</td>
<td>0.405</td>
<td>11.12</td>
<td>61</td>
</tr>
<tr>
<td></td>
<td>1.039</td>
<td>0.535 ± 0.059</td>
<td>0.403</td>
<td>11.12</td>
<td>71</td>
</tr>
<tr>
<td></td>
<td>1.245</td>
<td>0.619 ± 0.061</td>
<td>0.281</td>
<td>11.10</td>
<td>63</td>
</tr>
<tr>
<td></td>
<td>0.699</td>
<td>0.671 ± 0.051</td>
<td>0.941</td>
<td>10.30</td>
<td>85</td>
</tr>
<tr>
<td></td>
<td>0.807</td>
<td>0.570 ± 0.046</td>
<td>1.109</td>
<td>10.24</td>
<td>114</td>
</tr>
<tr>
<td>Low-$M_*/\text{high-log}(1+\delta_{\text{gal}})$</td>
<td>0.901</td>
<td>0.460 ± 0.053</td>
<td>1.049</td>
<td>10.28</td>
<td>87</td>
</tr>
<tr>
<td></td>
<td>1.030</td>
<td>0.309 ± 0.062</td>
<td>0.794</td>
<td>10.22</td>
<td>55</td>
</tr>
<tr>
<td></td>
<td>1.258</td>
<td>0.143 ± 0.059</td>
<td>0.702</td>
<td>10.24</td>
<td>35</td>
</tr>
<tr>
<td></td>
<td>0.697</td>
<td>0.667 ± 0.059</td>
<td>0.953</td>
<td>10.66</td>
<td>63</td>
</tr>
<tr>
<td></td>
<td>0.812</td>
<td>0.654 ± 0.046</td>
<td>1.173</td>
<td>10.60</td>
<td>107</td>
</tr>
<tr>
<td>Med-$M_*/\text{high-log}(1+\delta_{\text{gal}})$</td>
<td>0.902</td>
<td>0.507 ± 0.059</td>
<td>1.090</td>
<td>10.73</td>
<td>146</td>
</tr>
<tr>
<td></td>
<td>1.059</td>
<td>0.560 ± 0.050</td>
<td>0.875</td>
<td>10.66</td>
<td>100</td>
</tr>
<tr>
<td></td>
<td>1.257</td>
<td>0.455 ± 0.075</td>
<td>0.812</td>
<td>10.62</td>
<td>44</td>
</tr>
<tr>
<td></td>
<td>0.698</td>
<td>0.746 ± 0.057</td>
<td>0.958</td>
<td>11.09</td>
<td>59</td>
</tr>
<tr>
<td></td>
<td>0.804</td>
<td>0.690 ± 0.038</td>
<td>1.166</td>
<td>11.10</td>
<td>145</td>
</tr>
<tr>
<td>High-$M_*/\text{high-log}(1+\delta_{\text{gal}})$</td>
<td>0.874</td>
<td>0.644 ± 0.062</td>
<td>1.192</td>
<td>11.16</td>
<td>59</td>
</tr>
<tr>
<td></td>
<td>1.096</td>
<td>0.778 ± 0.049</td>
<td>0.829</td>
<td>11.08</td>
<td>72</td>
</tr>
<tr>
<td></td>
<td>1.262</td>
<td>0.621 ± 0.064</td>
<td>0.839</td>
<td>11.11</td>
<td>58</td>
</tr>
</tbody>
</table>
Figure 6. Average quiescent fraction ($f_q$) versus redshift for cluster galaxies in three different stellar mass bins simulated using the method described in Appendix A. Red diamonds, green squares, and blue circles show the average $f_q$ values and their associated NMAID uncertainties for 200 sets of simulated galaxies equivalent to the high-, med-, low-$M_*$ samples. In these simulations the effects of environmentally driven quenching are purposely not included. These effects are omitted in order to create a reference sample of group/cluster galaxies with which to compare to similar galaxies observed in ORELSE groups and clusters. As both galaxy populations are, by construction, imprinted with the integrated effect of secular processes, and the latter sample contains the integrated effect of environmental processes, this comparison can provide an estimate of the latter (see Fig. 7).

Figure 7. Quenching efficiency, defined as $f_q^{\text{high,log}(1+\delta_{\text{gal}})} - f_q^{\text{ref,log}(1+\delta_{\text{ref}})}$, as measured in a variety of different surveys for galaxy populations with different stellar mass ranges. The reference sample is, ideally, a sample of galaxies comparable to those galaxies observed in high-density environments but which have been largely untouched by physical effects specific to environment. Different schemes are used to define a reference sample. For the ORELSE measurements, shown along with their uncertainties by the red, green, and blue shaded regions, the reference sample was generated using the simulations presented in Section 3.2 and Appendix A. For other surveys, shown as grey points (diamonds: Nantas et al. 2017, circles: Fossati et al. 2017, inverted triangles: Balogh et al. 2016, triangle: Kovač et al. 2014, stars: Muzzin et al. 2012 and Balogh et al. 2016), the reference sample is typically chosen as set of field galaxies coeval to and stellar-mass matched to a set of high-density galaxies (see Section 3.2). The size of the points scale with the stellar mass of the sample, with small, medium, and large points roughly corresponding to the same stellar mass ranges as the low-, med-, and high-$M_*$ ORELSE samples, respectively.

et al. 2012) as reported in Balogh et al. (2016). Though not plotted, the $\Psi_{\text{conv}}$ values derived here are consistently lower at essentially all $M_*$ and redshifts than those measured for groups and cluster populations from initial data taken as part of the Hyper Suprime-Cam Subaru Strategic Program (Jian et al. 2018). However, the large number of variants on the method of calculation of this quantity and our imperfect knowledge of any differences that exist the galaxy samples presented in these works and those selected here do not allow us to claim anything other than the ORELSE measurements of environmental quenching efficiency appear broadly consistent with those from other surveys.

Though not plotted, the high value of the environmental quenching efficiency observed for low-$M_*$ galaxies in our lowest redshift bin is broadly consistent with the strong signatures of environmentally driven quenching seen in the VIPERS Multi-Lambda Survey (Moutard et al. 2016a) for similar galaxies at similar redshifts (Moutard et al. 2018). In that study, the authors constrain the quenching time-scales for such galaxies and settle on ‘delayed-then-rapid’ picture of quenching similar to that presented in Section 3.3 in our study. Further, it appears that such a quenching channel, while perhaps not the dominant channel for the full population of higher stellar mass galaxies at lower redshifts (Moutard et al. 2018), appears dominant for galaxies of these stellar masses which inhabit clusters and groups at all redshifts studied here (see Section 3.3 for more discussion). However, that study lacked the redshift baseline to observe the precipitous drop of environmental quenching efficiency for galaxies $\log(M_*/M_\odot) < 10.4$. These lines of evidence generally point to a picture in which the conditions in groups and clusters, while still sufficient to be able to efficiently quench the highest stellar mass galaxies over the redshift range studied here, are changing subtly with redshift. That the environmental quenching efficiency for lower-$M_*$ declines rapidly over this window perhaps suggests a similar decline for higher-$M_*$ at even higher redshifts than are studied here. Such a picture is consistent with several works studying clusters at $z \gtrsim 1.4$, in which the average star formation is seen in many cases to be unperturbed or even enhanced relative to field environments (e.g. Brodwin et al. 2013; Alberts et al. 2014, 2016; Santos et al. 2014; Wang et al. 2016; Stach et al. 2017; Coogan et al. 2018).

That efficient environmental quenching is observed to persist up to $z \lesssim 1.4$ for galaxies of stellar masses $\log(M_*/M_\odot) \gtrsim 10.4$ is consistent with results presented in a companion ORELSE study (Tomczak et al. 2019). In this study, the specific star formation rate– overdensity for star-forming galaxies and $f_q$–overdensity relations from a combined spectroscopic and photometric sample drawn from ORELSE are constructed and used to infer the relative number of galaxies undergoing active quenching in different environments in different stellar mass bins across the redshift range $0.8 \lesssim z \lesssim 1.05$. For galaxies in a similar stellar mass range, the majority of galaxies (70–90 per cent) inhabiting the highest density environments observed in ORELSE $\log (1 + \delta_{\text{gal}}) > 1$ are estimated to be undergoing quenching characterized by an $SFR$ that is exponentially declining with $\tau = 0.5$ Gyr. In the next section, we will combine these results with N-body simulations and the $\Psi_{\text{conv}}$ valued derived here to create estimates of various time-scales required to convert a galaxy from star-forming to quiescent after the accretion into a group/cluster environment. Interestingly, the clear persistence of the colour–density relation for higher-stellar mass galaxies across the entire redshift range studied here may help to explain the relatively high quiescent fractions found in distant cluster surveys which are predominantly sensitive to galaxies above $\log(M_*/M_\odot) \gtrsim 10.4$ (e.g. Newman et al. 2014; Cooke et al. 2012) as reported in Balogh et al. (2016).
et al. 2016; Lee-Brown et al. 2017; Strazzullo et al. 2019). This persistence is also broadly consistent with the results of van der Burg et al. (2013) who studied the quiescent fraction as a function of stellar mass for galaxies inhabiting 10 massive $z \sim 1$ clusters drawn from the GCLASS survey. In that study, the authors found that quiescent cluster galaxies significantly outnumbered their star-forming counterparts for all but the lowest stellar masses probed in their data [$\log(M_*/M_\odot) \leq 10.2$].

Another interesting consequence of the results presented here is that ram pressure stripping and harassment are precluded as the sole mechanisms for driving environmental quenching over the redshift range $0.55 \leq z \leq 1.4$. Because these mechanisms rely on stripping matter, typically atomic or molecular gas, in situ to group/cluster galaxies, everything else being equal, e.g. galaxy matter density profile, differential velocity, the density of the intragroup or intracluster medium (IGM/ICM) in the case of ram pressure stripping, galaxies with a shallower potential well are more susceptible to such processes (e.g. Gunn & Gott 1972; Fillingham et al. 2016). That $\Psi_{\text{convert}}$ is observed to decrease with redshift for the lowest stellar mass galaxies [i.e. $\log(M_*/M_\odot) \sim 10.2$] without a corresponding decrease for their higher stellar mass counterparts is a behaviour that cannot be reconciled with ram pressure stripping or harassment acting as the primary quenching mechanism unless lower stellar mass galaxies are accreted with a higher (total) amount of gas or are generally more concentrated than galaxies with higher stellar masses. Neither of these differences are observed in studies of field galaxies in the redshift range studied here (e.g. Popping et al. 2012; van der Wel et al. 2014; Scoville et al. 2017). Alternatively, if the conditions in groups and clusters change by $z \sim 1.4$ to reduce the efficacy of ram pressure stripping or harassment, i.e. the differential velocity of galaxies relative to one another or the density of the IGM/ICM decreases, this change would possibly provide some possibility of explaining the observed behaviour. However, such a change in conditions is extremely unlikely for the sample of ORELSE clusters as they are observed to contain a well-developed ICM and a galaxy population with a high $\sigma_v$.

In Fig. 8 we plot the scaled conversion time-scale from the measurements of $\Psi_{\text{convert}}$ and the simulated accretion history, as we are able to determine from the latter, the percentage of galaxies accreted by a given time, e.g. by $(z) = 0.65$, approximately 50 per cent of star-forming galaxies have been in the cluster environment for longer than $\gtrsim 3$ Gyr (see the lookback time corresponding to the point at which the fraction of halo mass accreted is equal to 0.5 in the top left panel of Fig. 8). Notice that we make the additional assumption here that the mass accreted in all time-steps of the Millennium simulation scales by a constant multiplicative factor with respect to the mass accreted in the form of star-forming galaxies. It is likely that this assumption that broadly holds as the $f_q$ values of accreted galaxies, i.e. those in the field of the stellar mass ranges considered here, are relatively flat for the majority of the accretion history of the cluster (see Fig. 6 and Tomczak et al. 2014). Further, we assume on average the amount of stellar mass loss or growth of each star-forming galaxy from the time of accretion until it becomes quiescent is insufficient to move into a different stellar mass bin. The wideness of the stellar mass bins used here help immensely with the validity of this assumption. Finally, as in the simulations presented in Appendix A the effects of merging processes in modulating the number or stellar mass distribution of galaxies since being accreted are again ignored here.
Persistence of environmental quenching to $z \sim 1.4$

Figure 8. Average halo mass accretion histories of two different types of structures differentiated by their $z = 0$ mass scaled as estimated from merger trees applied to the Millennium $N$-body simulation taken from McGee et al. (2009). The five panels correspond to measurements on simulated and real galaxies falling in the five different redshift ranges applied to the final ORELSE spectral sample, with the average redshift of each sample shown in the top right of each panel. Accretion histories are estimated at $z = 0$ and scaled to the five average redshift values by the dynamical time $(1 + z)^{-3/2}$. The shaded colour regions in each panel denote the $\pm 1\sigma$ range of environmental quenching efficiencies, $\Psi_{\text{convert}}$, of the three different stellar mass samples in each redshift bin. The average stellar mass of each sample in each redshift bin is given in the top right of each panel. The intersection between the thick solid line and each shaded region sets the conversion time-scale, $\tau_{\text{convert}}$, of each sample at each redshift. Notice that while $\tau_{\text{convert}}$ remains relatively constant for the med- and high-$M_*$ samples across all redshifts studied here, $\tau_{\text{convert}}$ is seen to consistently increase with increasing redshift for low-$M_*$ galaxies, and becomes consistent with being longer than the time since formation of the cluster by $\langle z \rangle = 1.27$. 

them in Table 4. For higher stellar mass galaxies, i.e. those in the med- and high-\(M_\star\) bins, \(t_{\text{convert}}\) is seen to largely hold or slightly decrease over the entire redshift range studied here, 0.55 \(\leq \frac{z}{1.4}\), varying from 1.6–3.0 Gyr across the various samples with an average of 2.4 \(\pm\) 0.3 Gyr. Conversely, the conversion time-scale for low-\(M_\star\) galaxies is observed to monotonically increase with increasing redshift, spanning from 2.7–3.8 Gyr over the same redshift range, with an average of 3.3 \(\pm\) 0.3 Gyr, again suggesting that cluster environments struggle to process lower stellar mass galaxies at \(z \gtrsim 1.1\).

These results are in broad agreement with the few studies that have attempted to constrain similar or identical time-scales within the redshift range probed by the ORELSE sample. Foltz et al. (2018) studied the photometric and spectroscopic galaxy population of 14 clusters at 0.85 \(\leq \frac{z}{1.65}\) drawn from the GCLASS and SpARCS cluster surveys. Using an approach which combines the colour distributions of background-subtracted member populations and accretion rates estimated from galactic simulations they constrained \(t_{\text{convert}}\) (called \(t_\text{p}\) in their nomenclature) of member galaxies with \(\log(M_\star/M_\odot) \geq 10.5\). For member galaxies with mean redshifts of \(\langle z \rangle = 1.04\) and 1.55 the conversion time-scales were estimated at \(t_{\text{convert}} = 1.50^{+0.08}_{-0.10} + 1.24^{+0.22}_{-0.20}\). These time-scales are consistent with the values derived for our high-\(M_\star\) sample at comparable redshifts, though somewhat in tension with those values derived from our med-\(M_\star\) sample. However, the stellar mass distribution of their final sample appears to be a compromise between that of the ORELSE med- and high-\(M_\star\) galaxies. In addition, their sample is comprised exclusively of cluster galaxies while the high-log(1 + \(\delta_{\text{gal}}\)) ORELSE sample contains galaxies residing in filaments, groups, and clusters. These differences likely ease this tension considerably. Balogh et al. (2016) studied the spectroscopic member populations in 0.8 \(\lesssim \frac{z}{1.2}\) groups and clusters from the Galaxy Environment Evolution Collaboration 2 (GEEC2; Balogh et al. 2011, 2014) and GCLASS surveys. Combining measurements made from coeval field and group/cluster populations along with results from the same simulations that are employed here, they found conversion time-scales (called \(t_\text{p}\) in their nomenclature) of \(t_{\text{convert}} = 2.4 \pm 0.6\) and 1.7 \(\pm\) 0.6 for group and cluster galaxies, respectively, in the stellar mass bin 10.3 \(\leq \log(M_\star/M_\odot) < 10.75\). These values are consistent with the \(t_{\text{convert}}\) values measured in ORELSE at equivalent redshifts and stellar masses. Interestingly, in both the group and cluster sample, though the former employing \(t_{\text{phot}}\)-selected member galaxies, they found evidence of \(t_{\text{convert}}\) abruptly increasing with decreasing stellar mass, with estimated values consistent with those derived from the low-\(M_\star\) sample. Estimates of \(t_{\text{convert}}\) derived from groups in 3D-HST for galaxies of stellar masses comparable to the med-\(M_\star\) are also found to be comparable to the estimates from our sample over the same redshift range \(t_{\text{convert}} \sim 2.5 \pm 0.5\) Gyr, Fossati et al. (2017). While considerable uncertainties exist in the conversion time-scales of a particular galaxy population, these estimates of \(t_{\text{convert}}\) along with our own point to a picture in which galaxies spend a considerable time in the group or cluster environment prior to completely ceasing star formation activity and fading to quiescence.

Such a picture is, however, seemingly difficult to reconcile with the large fractional population of galaxies known as ‘post-starburst’ or ‘K + A’ galaxies seen in high redshift groups and clusters (e.g. Dressler et al. 1999; Tran et al. 2003; Muzzin et al. 2014; Wu et al. 2014; Lemaux et al. 2017; Socolovsky et al. 2018). Such galaxies contain spectral features of colours that necessarily indicate a rapid cessation in their star formation activity and appear in numbers considerable relative to those of star-forming galaxies (Poggianti et al. 2009; Lemaux et al. 2017), strongly implying that rapid quenching is a common pathway to quiescence for such galaxies. Additionally, as mentioned in the previous section, we have observed in ORELSE a strong indication of a fractionally large transition population amongst group/cluster galaxies with stellar masses in excess of \(\log(M_\star/M_\odot) > 10.1\) at 0.85 \(\leq \frac{z}{1.0}\). This transition population is characterized by an exponentially declining SFH with \(\tau = 0.5\). Assuming this form of the SFH, a star-forming galaxy at this redshift would have colours consistent with quiescence between 0.8 and 1.4 Gyr after the inception of the initial quenching event, values considerably shorter than any of the conversion time-scales estimated here. Such a time-scale is nearly identical to the luminosity-weighted stellar age of the oldest K + A galaxies observed in the ORELSE survey \(\langle t_{\text{K+A}} \rangle = 1.00^{+0.34}_{-0.21}\) Gyr; Lemaux et al. 2017). The only possibility to reconcile the relatively rapid time-scale involved in moving a galaxy to quiescence once quenching has begun and the relatively large time-scale of \(t_{\text{convert}}\) is to introduce a third time-scale during which a galaxy accreted to a group or cluster environment remains relatively unaffected by quenching processes. We will refer to this time as \(t_{\text{delay}}\) and the time from the inception of the quenching event to the time that a galaxy exhibits colours classifying it as quiescent as \(t_{\text{quench}}\) such that \(t_{\text{quench}} + t_{\text{delay}} \equiv t_{\text{convert}}\).

Adopting \(t_{\text{quench}}\) as 1.1 \(\pm\) 0.3 Gyr from Tomczak et al. (2019) and solving for \(t_{\text{delay}}\) yields an average \(t_{\text{delay}} = 1.3 \pm\) 0.4 and 2.2 \(\pm\) 0.4 Gyr for the combined med- and high-\(M_\star\) and low-\(M_\star\) ORELSE samples, respectively, across all redshifts studied. These delay times are broadly consistent with similar time-scales measured from other surveys in a few comparable redshift and stellar mass ranges (e.g. Mok et al. 2014; Muzzin et al. 2014; Balogh et al. 2016). These values are also well in excess of the average time required for galaxies to move to their first pericentric passage after being accreted by their parent group or cluster at the median redshift of the ORELSE sample (Wetzel 2011; Wetzel et al. 2013). Such a long delay time is supported by recent evidence taken from SDSS and redMaPPer clusters at low- to intermediate-redshift that show appreciable populations of galaxies far from centre of clusters thought to have survived their first passage through the cluster to make it to their first apogee (i.e. ‘splashback’ galaxies, More et al. 2016; Baxter et al. 2017; Chang et al. 2018). The fact that the bulk of group/cluster galaxies remain largely unperturbed in their star-formation activity at least through their initial pericentric passage and that an additional \(\sim 0.5\)–1.5 Gyr is allowed to elapse prior to the inception of quenching appears to place most galaxies in the more rarefied regions of their parent groups/clusters when quenching is initiated. This inference strongly disfavours ram pressure stripping and galaxy harassment as the sole mechanisms responsible for quenching. While orbits can take many forms other than purely radial and pericentric passage can occur far from the group/cluster core, at least some fraction of newly accreted galaxies must pass through \(R_{500}\). The lack of a significant quenching signature in this population makes the possibility that these processes are the sole inducer of quenching activity remote.

Conversely, dynamical processes such as galaxy–galaxy merging events or other strong interactions are strongly favoured by such a scenario, as the \(t_{\text{delay}}\) time-scales estimated here appears to be directly scaled from comparable time-scales measured at \(z = 0\) by the dynamical time (Wetzel et al. 2013). That merging events might be primarily responsible for catalysing quenching qualitatively conforms with the difference in the average \(t_{\text{delay}}\) observed for lower and med- to high-\(M_\star\) galaxies in ORELSE. While the major merging rate at these redshifts is not strongly dependent on stellar
mass in the ranges of stellar masses studied here at least in the field (e.g. López-Sanjuan et al. 2013), field galaxies accreted into the group/cluster environment at higher stellar masses are on average more highly clustered than those at lower stellar mass (e.g. Marulli et al. 2013; Coil et al. 2017; Durbale et al. 2018). The fact that the bulk of the galaxies accreted as star forming survive in that state past their initial pericentric passage provides an intriguing possibility that tidal forces related to the overarching halo or interactions with in situ galaxies provide the necessary frustration to induce merging between infalling pairs or small groups of galaxies. A merging scenario to induce the beginnings of quenching activity is also broadly supported in the analysis of Lemaux et al. (2017) in a study of the K + A populations of two ORELSE LSSs. In that study, it was proposed that merging events in the outskirts of groups and clusters induce strong stellar and/or AGN feedback that begin the quenching process and this quenching is then finalized by the effects of ram pressure stripping as the remnant galaxy makes it way back to pericentric passage. In this scenario, ram pressure stripping is invoked only to prevent the galaxy from rejuvenating its star-formation activity rather than as a primary transformative mechanism. The time-scales estimated above appear to largely support such a scenario as the additional quenching time, \( t_{\text{quench}} \), is long enough to allow galaxies the possibility of making their second passage through their parent group or cluster.

This scenario partially mirrors that put forth in Brodwin et al. (2013), in which quenching of galaxies inhabiting higher redshift clusters \((z \gtrsim 1)\) was suggested to proceed through a combination of major merging and AGN activity. Such a scenario is supported by the increase in both types of activity with increasing redshift (see e.g. Mancone et al. 2010; Tomczak et al. 2017; Marulli et al. 2013), in which quenching of galaxies inhabiting higher redshift clusters \((z \gtrsim 1.15)\) was suggested to proceed through a combination of major merging and AGN activity. Such a scenario is supported by the increase in both types of activity with increasing redshift. While a large fraction of galaxies in the stellar mass \(M_\ast \gtrsim 10^{10} M_\odot\) in high-redshift clusters and groups induce strong stellar and/or AGN feedback that begin the quenching process and this quenching is then finalized by the effects of ram pressure stripping as the remnant galaxy makes it way back to pericentric passage. In this scenario, ram pressure stripping is invoked only to prevent the galaxy from rejuvenating its star-formation activity rather than as a primary transformative mechanism.

The above ‘merging-then-stripping’ scenario appears to be broadly able to explain the relatively high quenching efficiencies observed at all redshifts for the med- and high-\(M_\ast\) samples. However, it is not clear such a scenario is appropriate for low-\(M_\ast\) galaxies. At \(z \sim 0.7\), \(\Psi_{\text{conv}}\) for low-\(M_\ast\) galaxies is indistinguishable from that of their higher-\(M_\ast\) counterparts. But, unlike higher-\(M_\ast\), populations, environmental quenching appears to become less effective with increasing redshift for low-\(M_\ast\) galaxies and is completely ineffectual at the highest redshifts studied here.

Given this dramatically different behaviour, it naively appears that the merging-then-stripping scenario proposed above cannot explain the behaviour of both sets of galaxies.

However, it is possible that various considerations related to the scenario above that allow for efficient quenching of higher-\(M_\ast\) galaxies at all redshifts operate on low-\(M_\ast\) populations in a different way. Further, it is possible that these differences act in such a way to explain the decreased quenching efficiency of these galaxies in high-\(z\) clusters and groups. The main considerations within the scenario are (a) the likelihood of undergoing a wet or mixed merging event, (b) subsequent to a merging event, the likelihood of sufficient stellar and/or AGN feedback to push gas that survives to large galactocentric distance, and (c) the likelihood that a parent cluster or group has a sufficiently hot and dense medium to strip the excised gas from the merger remnant. If any of these three considerations are not the same for low-\(M_\ast\) as for higher-\(M_\ast\) galaxies, it is possible that the merging-then-stripping scenario can still be reconciled with the behaviour observed for galaxies of all stellar masses considered in this study.

Taking each in turn, the merging rate of field galaxies is observed to either stay constant or increase with increasing redshift over the redshift range studied here (Bundy et al. 2009; López-Sanjuan et al. 2013) and does not depend strongly on stellar mass implying (a) is equally or more likely to happen at high redshift. Thus, it is unlikely that (a) is appreciably different for low-\(M_\ast\) galaxies at higher \(z\) than for other samples. However, the gas fraction of the galaxies comprising those mergers over that same redshift range increases by a factor of \(\sim 3\) for all stellar masses considered here (Scoville et al. 2017; Darvish et al. 2018) and becomes the dominant component of the baryonic content at least for low-\(M_\ast\) galaxies. As such, the feedback required under (b) is necessarily stronger than that at lower redshift. While stellar feedback is capable of providing some of the required effect, its effect alone is typically insufficient to excise the bulk of the gas to large galactocentric distances for all of the galaxies studied here [i.e. log(\(M_\ast/\odot\)] \(\gtrsim 10\); Bahé & McCarthy 2015, though see also Diamond-Stanic et al. 2012]. Activity from an AGN can substantially increase these large-scale outflows (e.g. Hopkins et al. 2016). While a large fraction of galaxies in the stellar mass range of the low-\(M_\ast\) galaxies studied here are known to host AGNs, these AGNs are typically of the less powerful variety (Juneau et al. 2013). As a consequence, it is possible that such AGN activity, in tandem with stellar feedback, is sufficient to provide the required feedback in such galaxies at the lower redshift end of our sample, but fails for the larger gas reservoirs present at higher redshifts. Incidentally, higher stellar mass galaxies would likely not suffer from the same considerations, as AGN activity, when present, is more likely to be powerful (e.g. Kauﬀmann et al. 2003; Pimbblet et al. 2013; Bongiorno et al. 2016; Lanzuisi et al. 2017).

Finally, we consider concerns related to (c). As we have mentioned previously, the median mass of the detected ORELSE structures does not change appreciably with redshift, nor do the properties of ICM of those ORELSE clusters detected at X-ray.
wavelengths (Rumpl et al. 2018), implying that (c) is equally likely at low and high redshift within the sample studied here. However, ORELSE groups are also observed to contain a larger fraction of lower stellar mass star-forming galaxies relative to cluster environments (see e.g. Tomczak et al. 2017, where local overdensity is used as a proxy for group and cluster environments). It is, thus, possible that low-\(M_\ast\) galaxies are generally experiencing less compelling ram pressure stripping than their higher-\(M_\ast\) counterparts. This logic, though, applies to low-\(M_\ast\) galaxies across all epochs studied here unless groups at lower redshifts, simply by virtue of having a longer available time to form, have been able to produce a denser medium (see e.g. Jeltema et al. 2009 for a discussion on this effect over a larger temporal baseline).

It, thus, appears that the conditions governing (b) and (c) are sufficiently uncertain to leave open the possibility that the merging-then-stripping scenario advocated for quenching group/cluster med- and high-\(M_\ast\) galaxies could also generally apply to low-\(M_\ast\) galaxies in those same environments without inducing similar levels of quenching at the highest redshifts probed in this study. The implementation of a variety of AGN selection techniques and obtaining deeper X-ray observations, both to uncover weaker AGN activity and to detect the diffuse medium of lower mass structures, as well as obtaining imaging from the Atacama Large Millimeter/submillimeter Array (Wootten & Thompson 2009) to measure the gas content of galaxy samples over a variety of stellar masses, environments, and redshifts would be extremely helpful to validate or disprove this scenario. Additionally, while our current detection techniques do not allow us to identify groups in a systematic and well-characterized manner up to the highest redshifts studied here, future work (Hung et al. 2019) using more sophisticated group and cluster detection algorithms applied to our VMC maps will allow the systematic detection of structures for \(\log(M_{\text{tot}}/M_\odot) > 13.5\) over the redshift range \(0.55 \leq z \leq 1.4\). Once these environments are identified, we will be able to additionally bin our stellar mass, environmental, and redshift sub-samples by total structure mass to further test this scenario.

4 SUMMARY AND CONCLUSIONS

In this study we presented the fraction of different types of galaxies which have ceased to form stars as estimated from observations taken as part of the ORELSE survey. These populations were drawn from a sample of \(\sim 5000\) spectroscopically confirmed galaxies which span a large range in stellar mass, environments, and redshifts and were shown to be broadly representative of the overall galaxy populations at these epochs. Using these fractions, in conjunction with a variety of models, we were able to place constraints on the average efficacy of environmental quenching over a large baseline in cosmic time for a large variety of galaxy populations, as well as time-scales and processes associated with this quenching. Our main conclusions are as follows:

(i) The quiescent fraction, \(f_q\), was generally observed to increase with increasing galaxy density, increasing stellar mass, and decreasing redshift for the range of environmental densities \(0.1 \leq \log(1 + \delta_{\text{gal}}) \leq 2.0\), stellar masses \(10 \leq \log(M_\ast/M_\odot) \leq 11.6\), and redshifts \(0.55 \leq z \leq 1.4\) studied here.

(ii) Except for the very lowest stellar mass bin at the highest redshifts probed, the colour-density relation was observed to persist to \(z \sim 1.4\). For intermediate and higher stellar mass galaxies \(\log(M_\ast/M_\odot) > 10.4\), essentially no significant decrease in \(f_q\) was observed from \(z = 0.55\) to \(z = 1.4\) at fixed stellar mass or environment. For lower stellar mass galaxies, the colour-density relation appeared to evolve rapidly with redshift over the redshift range studied here, essentially disappearing at the highest redshifts.

(iii) A general formula to calculate the quiescent fraction of galaxy populations was estimated and took the form:

\[
\Psi(\alpha, \beta, \gamma, \theta, \zeta) = \alpha z^\gamma + \beta \log(M_\ast/M_\odot)^\theta + \gamma \log(1 + \delta_{\text{gal}})^\zeta,
\]

with \(\alpha = -1.8544 \pm 0.0002\), \(\beta = 8.3876 \pm 0.0001\), \(\gamma = -11.8321 \pm 0.0002\), \(\theta = 0.1421 \pm 0.0007\), \(\zeta = 0.2224 \pm 0.0007\), and \(\kappa = -0.0076 \pm 0.0005\). This formula was designed to be applicable to general galaxy populations within the range of environments, stellar masses, and redshifts. However, it is possible that this formula slightly overestimates the quiescent fraction of lower stellar mass galaxies \(10 \leq \log(M_\ast/M_\odot) \leq 10.4\) by up to \(\sim 0.1\).

(iv) A semi-empirical model was used to generate a variety of mock galaxies that inhabit dense environments but which were shielded from the effects of environmental quenching. These predicted \(f_q\) values, in conjunction with the observed \(f_q\) values in dense environments, were used to estimate environmental quenching efficiencies (\(\Psi_{\text{convert}}\)) for all galaxy types. For intermediate and higher stellar mass galaxies, the efficiency of cluster/group environments to inhibit star-formation activity in such galaxies remains roughly constant for all redshifts studied here. For lower stellar mass galaxies, a precipitous drop in \(\Psi_{\text{convert}}\) was observed from lower to higher redshift.

(v) These \(\Psi_{\text{convert}}\) values along with group/cluster accretion histories as estimated from the Millennium simulation allowed constraints to be placed on the average time elapsed between the accretion of a galaxy and quiescence (\(t_{\text{convert}}\)). For intermediate and higher stellar mass galaxies, this value was \(t_{\text{convert}} = 2.4 \pm 0.3\), while for lower stellar mass galaxies this time was estimated to be significantly longer on average \((t_{\text{convert}}) = 3.3 \pm 0.3\).

(vi) Using results from a companion ORELSE study, which investigated the relationship between density and specific star-formation rate for a nearly identical set of galaxies (Tomczak et al. 2019), and the \(t_{\text{convert}}\) time-scales estimated here, we found strong support for a scenario in which a galaxy spends a considerable amount of time in a group/cluster environment unquenched followed by a rapid truncation of its star-formation, a phenomena known as ‘delayed-then-rapid’ quenching. These results were used to place constraints on the average delay time between accretion and the inception of the quenching process, finding \(t_{\text{delay}} = 1.3 \pm 0.4\) and \(t_{\text{delay}} = 2.2 \pm 0.4\) Gyr for the combined intermediate- to high-stellar mass and lower-stellar mass samples, respectively.

(vii) The behaviour of \(f_q\) with redshift as well as the long delay times estimated for all galaxy populations were both used to argue against ram pressure stripping and galaxy harassment as primary quenching mechanisms. Galaxies appear to make their passage through the group/cluster essentially unperturbed and the inception of the quenching event appears to occur in the more rarefied regions of the group/cluster environment. These results are, rather, consistent with a scenario of galaxies merging in the outskirts of clusters and groups prior to their second pericentric passage at least for galaxies \(\log(M_\ast/M_\odot) > 10.45\). It is possible that ram pressure stripping and/or cluster tidal effects aid the quenching process as the merger remnant returns to pericentre, stripping any remaining gas ejected to large galacto-centric distances. It appears, however, that these mechanisms cannot solely be responsible for the observed behaviour.
Persistence of environmental quenching to $z \sim 1.4$

While the redshift baseline of this study is large and extends to a lookback time of $\sim 9$ Gyr, it appears that the processes that serve to quench galaxies in higher density environments are still broadly efficient for the bulk of the massive galaxy population. However, the mere existence of a large population of relatively massive quiescent galaxies in the ORELSE cluster/groups across all redshifts studied here, their relative scarcity at higher redshift, and the presence of extremely massive star-forming galaxies in all redshifts studied here, strongly argues that the efficiency of environmental quenching must subside quickly, and perhaps even reverse its effects, in forming groups and clusters at earlier epochs ($z \gtrsim 1.4$). Studies of such forming structures will allow for investigations on the progenitor galaxies of those massive galaxies observed to be pervasively quenched at $z < 1.4$. Making statistical connections between these forming structures and more established structures at lower redshift, as well as their constituent galaxy populations, will allow for a comprehensive view of the average life of galaxies in dense environments from inception until quiescence.

ACKNOWLEDGEMENTS

This material is based upon work supported by the National Science Foundation under Grant No. 1411943. Part of the work presented herein is supported by NASA Grant Number NNX15AK92G. This work was additionally supported by the France-Beckery Foundation, a joint venture between UC Berkeley, UC Davis, and le Centre National de la Recherche Scientifique de France promoting lasting institutional and intellectual cooperation between France and the United States. PFW acknowledges funding through the H2020 ERC Consolidator Grant 683184 and the support of an EACOA Fellowship from the East Asian Core Observatories Association. BCL gratefully acknowledges Julie Nantais, Susmita Adhikari, and Olga Cucciati for discussions that helped improve the paper. This paper is dedicated to Shari and Gary Goodenow, whose generosity and humor had a deep effect on the formative years of the first author, and whose influence and love is still profoundly felt. This study is based, in part, on data collected at the Subaru Telescope and obtained from the SMOKA, which is operated by the Astronomy Data Center, National Astronomical Observatory of Japan. This work is based, in part, on observations made with the Spitzer Space Telescope, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. UKIRT is supported by NASA and operated under an agreement among the University of Hawaii, the University of Arizona, and Lockheed Martin Advanced Technology Center; operations are enabled through the cooperation of the East Asian Observatory. When the data reported here were acquired, UKIRT was operated by the Joint Astronomy Centre on behalf of the Science and Technology Facilities Council of the U.K. This study is also based, in part, on observations obtained with WIRCAM, a joint project of CFHT, Taiwan, Korea, Canada, France, and the Canada–France–Hawaii Telescope which is operated by the National Research Council (NRC) of Canada, the Institut National des Sciences de l’Univers of the Centre National de la Recherche Scientifique of France, and the University of Hawai‘i. Some portion of the spectrographic data presented herein was based on observations obtained with the European Southern Observatory Very Large Telescope, Paranal, Chile, under Large Programs 070.A-9007 and 177.A-0837. The remainder of the spectrographic data presented herein were obtained at the W.M. Keck Observatory, which is operated as a scientific partnership among the California Institute of Technology, the University of California, and the National Aeronautics and Space Administration. The Observatory was made possible by the generous financial support of the W.M. Keck Foundation. We thank the indigenous Hawaiian community for allowing us to be guests on their sacred mountain, a privilege, without which, this work would not have been possible. We are most fortunate to be able to conduct observations from this site.

REFERENCES

Cucciati O. et al., 2010, A&A, 524, A2

Persistence of environmental quenching to $z \sim 1.4$

When calculating environmental quenching efficiencies, it is necessary to define a sample of galaxies that has had the chance to experience environmental quenching, if such quenching exists, and a sample of galaxies that has evolved on a nearly identical track, but has been completely sheltered from processes related to environmental quenching. Defining the former set of galaxies becomes a trivial task with the data available to the ORELSE survey as large number of galaxies have been spectroscopically confirmed in group and cluster environments. The sample that was adopted for this part of the calculation was referred to as the high-log (1 + $\delta_{gal}$) sample in the main body of the paper. The definition of the latter population is more tricky. In many studies, the population that has evolved on a similar track as group/cluster galaxies, though without the possibility of experiencing pervasive environmental quenching, is selected as coeval field populations. While this set of galaxies can serve as an approximation of what group/cluster galaxies would have evolved into in the absence of environmental quenching, there are several issues that make this approximation a crude one. The most pernicious of these issues is that field galaxies at a given redshift cannot be the progenitors of and group/cluster galaxies at that same redshift, and it is, rather, field galaxies at earlier epochs that seed the group/cluster population at a given redshift. Further, it is not possible to fully expiate this issue by substituting a field galaxy population observed at a fixed epoch earlier than that of the group/cluster population as group/cluster progenitor galaxies originate from a spectrum of earlier epochs. Other issues include the differential growth of stellar mass between field and group/cluster populations and differing selections of galaxies in field versus group/cluster environments.

In an attempt to circumvent these issues, we generated a suite of mock galaxies that were intended to mimic the properties of group and cluster galaxies in absence of all environmental quenching processes. As such, in this simulation we completely ignored the effects of merging, ram-pressure stripping, tidal stripping, harassment, strangulation, and all other galaxy–galaxy and galaxy-group/cluster processes that might effect the star formation and stellar mass properties in excess of what is experienced in the field.

Because we are primarily focused on the quiescent fraction across relatively large ($\sim 0.5$ dex) bins in stellar mass, we assume for the
presents the amount of accreted into the (proto-)cluster environment 
by the total mass at \( z = 0 \). This value sets the amount of mass accreted into the (proto-)cluster environment in this time-step. In order to translate this amount of halo mass into a galaxy population, we begin sampling from the galaxy total SMF presented in Tomczak et al. (2014) appropriate for the median redshift of the step in this case, \( z \approx 3.85 \). For each galaxy sampled in this way, a galaxy was assigned:

(i) A stellar mass, \( M_\star \).
(ii) A halo mass, \( M_{\text{halo}} \), estimated from inverting the redshift-dependent halo-to-stellar mass relation of Moster et al. (2010, 2013).
(iii) An initial state of quiescence or ongoing star formation, a state that is chosen from sampling a random uniform variable and comparing that sample to the fraction of quiescent to star-forming galaxies in the field at that redshift and \( M_\star \), as estimated by Tomczak et al. (2014) (e.g. if the random sample for a simulated galaxy was 0.61 and the fraction of star-forming to total galaxies at that redshift and \( M_\star \) was 0.6, the galaxy is considered quiescent).
(iv) For those galaxies assigned a state of ongoing star formation, a star formation rate (SFR) was assigned based on the redshift-dependent SFR–\( M_\star \) relation for star-forming galaxies of Tomczak et al. (2016) with an imposed scatter of 0.15 dex.

(v) An exponentially declining star formation future with a normalization set by the assigned SFR at the time of accretion and an e-folding time equal to 2 Gyr, a value appropriate for the general field population at the stellar masses of interest here (e.g. Tomczak et al. 2019).

Only those galaxies with stellar masses in excess of \( \log(M_\star/M_\odot) \geq 9 \) are considered in the simulation. This limit is roughly the stellar mass limit of the SMFs presented in Tomczak et al. (2014) at all of the redshifts considered here and is sufficient to generate essentially all galaxies in the stellar mass and redshift ranges of interest. For this first step, the SMF continues to be sampled until the amount of halo mass assembled reaches or exceeds 95 per cent of the halo mass predicted to be accreted in that time-step, a threshold that typically requires \( \sim 10 \) galaxies to be generated. At the conclusion of this first step, all mock galaxies are binned into the same three stellar mass bins considered throughout this paper, and \( f_a \) is calculated for each bin.

In subsequent steps, two galaxy populations are considered: newly accreted galaxies and those galaxies that are already coalesced into the cluster environment in previous steps. For the former population, the same steps as listed above are repeated and these galaxies are incorporated into the full simulated cluster galaxy population. For the latter population, the amount of stellar mass generated through in situ star formation in that step is added to the previous stellar mass of the galaxy by integrating star formation rate over that step and subtracting the amount of stellar mass lost over that same time period resulting from stellar evolution (BC03, Chabrier 2003; Moster et al. 2013). The value of \( f_a \) is then calculated for the combined sample of galaxies in the three \( M_\star \) bins considered throughout this paper.

At the conclusion of this process, the SFR of each galaxy is evolved by the time elapsed in the step subject to its star formation future set at the moment of coalescence. Finally, we compute the specific star formation rate (SSFR) of all galaxies at the end of this time-step and set the star formation state of any galaxies with a log[SSFR/yr\(^{-1}\)] \( \leq -11 \) to quiescent. This SSFR threshold is chosen as it has been shown that galaxies below this limit are almost exclusively found in the NUVrJ quiescent region (Ilbert et al. 2013; Cucciati et al. 2017). A similar exercise is not performed for accreted field galaxies as those galaxies which have their initial state set as star forming, their SSFRs are set by the star-forming SFR–\( M_\star \) relation, and, thus, by construction, cannot be considered quiescent. Once quiescent, the simulation ignores the potential for rejuvenation (e.g. Kannappan, Guie & Baker 2009; Kannappan et al. 2013) of the star formation for each galaxy for the remainder of the time of the simulation. The results of these simulations are shown in the main body of the paper in Fig. 6.