



HAL
open science

NuSTAR spectral analysis of two bright Seyfert 1 galaxies: MCG +8-11-11 and NGC 6814

A. Tortosa, S. Bianchi, A. Marinucci, G. Matt, R. Middei, E. Piconcelli, L. W. Brenneman, M. Cappi, M. Dadina, A. de Rosa, et al.

► **To cite this version:**

A. Tortosa, S. Bianchi, A. Marinucci, G. Matt, R. Middei, et al.. NuSTAR spectral analysis of two bright Seyfert 1 galaxies: MCG +8-11-11 and NGC 6814. *Monthly Notices of the Royal Astronomical Society*, 2018, 473 (3), pp.3104-3112. 10.1093/mnras/stx2457. hal-03014218

HAL Id: hal-03014218

<https://hal.science/hal-03014218v1>

Submitted on 22 Nov 2024

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.



Distributed under a Creative Commons Attribution 4.0 International License

NuSTAR spectral analysis of two bright Seyfert 1 galaxies: MCG +8-11-11 and NGC 6814

A. Tortosa,^{1*} S. Bianchi,¹ A. Marinucci,¹ G. Matt,¹ R. Middei,¹ E. Piconcelli,²
L. W. Brenneman,³ M. Cappi,⁴ M. Dadina,⁴ A. De Rosa,⁵ P. O. Petrucci,⁶ F. Ursini⁴
and D. J. Walton⁷

¹*Dipartimento di Matematica e Fisica, Università degli Studi Roma Tre, via della Vasca Navale 84, I-00146 Roma, Italy*

²*Osservatorio Astronomico di Roma (INAF), via Frascati 33, I-00040 Monte Porzio Catone (Roma), Italy*

³*Smithsonian Astrophysical Observatory, 60 Garden St, MS-4, Cambridge, MA 02138, USA*

⁴*INAF-IASF Bologna, Via Gobetti 101, I-40129 Bologna, Italy*

⁵*INAF/Istituto di Astrofisica e Planetologia Spaziali, via Fosso del Cavaliere, I-00133 Roma, Italy*

⁶*Univ. Grenoble Alpes, CNRS, IPAG, F-38000 Grenoble, France*

⁷*Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge CB3 0HA, UK*

Accepted 2017 September 20. Received 2017 September 20; in original form 2017 July 14

ABSTRACT

We report on the *NuSTAR* observations of two bright Seyfert 1 galaxies, namely MCG +8-11-11 (100 ks) and NGC 6814 (150 ks). The main goal of these observations was to investigate the Comptonization mechanisms acting in the innermost regions of an active galactic nucleus (AGN) which are believed to be responsible for the UV/X-ray emission. The spectroscopic analysis of the *NuSTAR* spectra of these two sources revealed that although they had different properties overall (black hole masses, luminosity and Eddington ratios), they had very similar coronal properties. Both presented a power-law spectrum with a high-energy cut-off at $\sim 150\text{--}200$ keV, a relativistically broadened Fe K α line and the associated disc reflection component, plus a narrow iron line likely emitted in Compton thin and distant matter. The intrinsic continuum was well described by Comptonization models that show for MCG +8-11-11 a temperature of the coronal plasma of $kT_e \sim 60$ keV and an extrapolated optical depth $\tau = 1.8$; for NGC 6814, the coronal temperature was $kT_e \sim 45$ keV with an extrapolated optical depth of $\tau = 2.5$. We compare and discuss these values to some most common Comptonization models that aim at explaining the energy production and stability of coronae in AGNs.

Key words: galaxies: active – galaxies: individual: MCG +8-11-11 – galaxies: individual: NGC 6814 – galaxies: Seyfert – X-rays: galaxies.

1 INTRODUCTION

The primary X-ray emission of active galactic nuclei (AGNs), according to the standard paradigm, is due to thermal Comptonization of the soft disc photons in a hot, optically thin plasma (the so-called corona) located above the accretion disc (e.g. Shapiro, Lightman & Eardley 1976; Sunyaev & Titarchuk 1980; Haardt & Maraschi 1991, 1993; Haardt & Matt 1993; Haardt, Maraschi & Ghisellini 1994). The spectral shape of this component is, in the first approximation, a power law with a cut-off at high energy. The exact shape of the cut-off depends on the properties of the corona; it is customary to approximate it with an exponential function, at least for a preliminary, phenomenological analysis. The cut-off energy is

then expected to be two to three times the plasma electron temperature, while the power-law index is related to both the temperature and the optical depth of the corona (Petrucci et al. 2000, 2001). The primary emission can illuminate the accretion disc itself or more distant material, like the molecular torus, producing a reflection spectrum (George & Fabian 1991; Matt, Perola & Piro 1991). This spectrum is characterized by fluorescent lines, the most prominent being the Fe K α line at 6.4 keV (e.g. Nandra & Pounds 1994; Matt, Fabian & Reynolds 1997; Reynolds & Nowak 2003), and by a broad continuum peaking at $\sim 20\text{--}30$ keV, known as the Compton hump. When the reflection arises from the inner regions of the accretion disc, it is affected by the strong gravitational field of the black hole. The result is an Fe K α emission line with the shape modified into a skewed and asymmetric profile (Fabian et al. 1989, 2002, 2009; Tanaka et al. 1995). Since the relativistic lines originate within a few gravitational radii from the central object, the analysis of these

* E-mail: tortosa@fis.uniroma3.it

features represents a powerful probe of the innermost region of the AGN.

With the superior sensitivity of *NuSTAR* (*Nuclear Spectroscopic Telescope Array*; Harrison et al. 2013) above 10 keV, it is possible to separate the primary and reflected continua and measure the coronal parameters, breaking the degeneracy occurring when the high-energy cut-off cannot be measured. Indeed, a number of high-energy cut-off measurements in local Seyfert galaxies, on a wide range of Eddington ratios, have been already obtained (Fabian et al. 2015, 2017; Marinucci et al. 2016 and references therein) to investigate the Comptonization mechanisms acting in the innermost regions of an AGN and which are believed to be responsible for the X-ray emission. However more, and more precise measurements, are needed in order to put these studies on a more firm statistical ground and constrain the coronal parameters. NGC 6814 and MCG +8-11-11 are two bright radio-quiet unobscured Seyfert 1 galaxies; they are ideal sources for this goal.

MCG +8-11-11 ($z = 0.0204$) is a very X-ray bright AGN with a black hole mass of $\log \frac{M_{\text{BH}}}{M_{\odot}} = 7.19 \pm 0.02$ (Bian & Zhao 2003) and X-ray fluxes measured by *INTEGRAL*, of $F_{20-100 \text{ keV}} = 8.46 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ and $F_{2-10 \text{ keV}} = 5.62 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ (Malizia et al. 2012). The 2–10 keV absorption-corrected luminosity of the source is $6.45 \pm 0.04 \times 10^{43} \text{ erg s}^{-1}$ (Bianchi et al. 2010). *ASCA* (Grandi et al. 1998) and *BeppoSAX* (Perola et al. 2000) showed that the spectrum is well fitted by a model composed by a power law, a warm absorber, a Compton reflection component and an FeK α line. The best fit of the *ASCA* and *OSSE* data was an absorbed power law with spectral index $\Gamma = 1.73 \pm 0.06$ and an exponential cut-off at ~ 250 keV, plus a reflection component and a cold iron line. Also *BeppoSAX* data showed the presence of a cut-off at high energy (~ 170 keV). The *XMM-Newton* spectrum (Matt et al. 2006) revealed the lack of a soft excess, a large reflection component and a narrow iron line with a low equivalent width (EW) and no relativistic features. Bianchi et al. (2010) found in the *Suzaku* observation a relativistic Fe K α line, plus a narrow component with no associated reflection continuum.

NGC 6814 ($z = 0.0052$; Molina et al. 2009) is a Seyfert 1 Galaxy with black hole mass of $\log \frac{M_{\text{BH}}}{M_{\odot}} = 6.99^{+0.32}_{-0.25}$ (Pancoast et al. 2014, 2015) known to show X-ray variability by at least a factor of 10 over time-scales of years (Mukai et al. 2003). It is part of the reverberation mapping campaign ‘the LAMP project’ (Lick AGN Monitoring Project; Bentz et al. 2009). The hard and soft X-ray flux of this source is $F_{20-100 \text{ keV}} = 5.66 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ and $F_{2-10 \text{ keV}} = 0.17 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ (Malizia et al. 2012). The *INTEGRAL* spectrum (Malizia et al. 2014) showed that the source has a quite flat spectrum ($\Gamma = 1.68 \pm 0.02$) with an exponential cut-off at $E_c = 190^{+185}_{-66} \text{ keV}$. From the *XMM-Newton* observation, it is possible to see the presence of a narrow FeK α line (Ricci et al. 2014) with $\text{EW} = 82^{+17}_{-15} \text{ eV}$. The *Suzaku* observation shows significant variability, a primary continuum with a photon index of 1.53 ± 0.02 and no evidence of soft excess. It also shows the emission of the Fe K α and Fe XXVI lines with centroid energy and EW, respectively, $E_{\text{Fe K}\alpha} = 6.40 \pm 0.03 \text{ keV}$, $\text{EW}_{\text{Fe K}\alpha} = 170^{+30}_{-40} \text{ eV}$ and $E_{\text{FeXXVI}} = 6.94 \pm 0.07 \text{ keV}$, $\text{EW}_{\text{FeXXVI}} = 90^{+30}_{-40} \text{ eV}$ (Walton et al. 2013).

In this paper, we present the *NuSTAR* observations of MCG +8-11-11 (100 ks) and of NGC 6814 (150 ks), taken almost simultaneously with short *Swift* observations (20 ks each). In Section 2, the observations and data reduction are presented. In Section 3, we report on the spectral analysis of the two sources. The results are discussed and summarized in Section 4.

2 OBSERVATIONS & DATA REDUCTION

2.1 *NuSTAR*

MCG +8-11-11 and NGC 6814 were observed by *NuSTAR* with its two coaligned X-ray telescopes (Focal Plane Modules A and B), respectively, on 2016 August 19, and on 2016 July 04. No other sources apart from the targets are apparent in the images. The *NuSTAR* data were reduced with the *NuSTAR* Data Analysis Software (*NuSTARDAS*) package (v. 1.6.0). Cleaned event files (level 2 data products) were produced and calibrated using standard filtering criteria with the *NUPIPELINE* task using the last calibration files available from the *NuSTAR* calibration data base (CALDB 20170120). The extraction radii of the circular region were 0.5 arcmin for source and 1.5 arcmin for background spectra; there are no other bright X-ray sources within 1.5 arcmin from MCG 8-11-11 and NGC 6814 and no other sources were present in the background region. Net exposure time, after this process, is 98 ks for MCG 8-11-11 and 148 ks for NGC 6814 for both FPMA and B. The spectra were binned in order to oversample the instrumental resolution by at least a factor of 2.5 and to have a signal-to-noise ratio (SNR) greater than 5 for both sources in each spectral channel.

2.2 *Swift*-XRT

MCG +8-11-11 and NGC 6814 were observed by *Swift* UVOT+XRT almost simultaneously with *NuSTAR* for a total exposure time of 20 ks each. *Swift* XRT spectra were extracted using the *XSELECT* (v2.4c) command line interface to the *FTOOLS* (Blackburn 1995).

If there is a pile-up, the measured rate of the source is high (above about 0.6 counts s^{-1} in the photon-counting mode). The easiest way to avoid problems related to the pile-up is to extract spectra using an annular region, thus eliminating the counts in the bright core, where the pile-up will occur.

The *Swift* XRT spectra resulted to have a high pile-up degree. We tested different annular extraction regions for the source, gradually increasing the inner radius. Even with large inner extraction radii, the pile-up was not removed and, moreover, the SNR became very low. We, therefore, decided not to use the *Swift* XRT data for the spectroscopic analysis.

3 DATA ANALYSIS

The spectral analysis has been performed with the *XSPEC* 12.9.0 software package (Arnaud 1996). Throughout the paper, errors correspond to 90 per cent confidence level for one interesting parameter ($\Delta\chi^2 = 2.7$), if not stated otherwise. The cosmological parameters $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_{\Lambda} = 0.73$ and $\Omega_m = 0.27$ are adopted.

Both sources show variability in their light curves (especially NGC 6814, well known to be a variable source; Mukai et al. 2003; Walton et al. 2013). The variability of NGC 6814 was consistent with the softer-when-brighter behaviour and with the fact that its black hole mass is 1/10 times the black hole mass of MCG +8-11-11 but since no strong spectral variations were found in the ratio between the 10–80 and 3–10 keV count rates (see Fig. 1), we decided to use time-averaged spectra for both sources.

We performed our data analysis by fitting the 3–80 keV *NuSTAR* spectra with different models, each of them including Galactic absorption with column densities $N_{\text{H}} = 1.84 \times 10^{21} \text{ cm}^{-2}$ for MCG +8-11-11 and $N_{\text{H}} = 9.11 \times 10^{20} \text{ cm}^{-2}$ for NGC 6814, as derived from H I maps (Kalberla et al. 2005). We tested the presence of

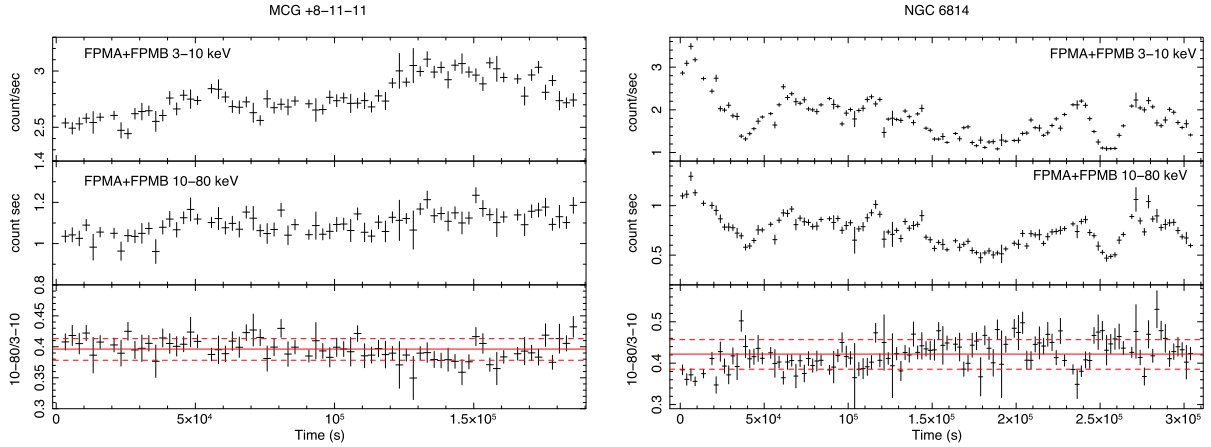


Figure 1. The *NuSTAR* FPMA+B light curves in the 3–10 keV (top panels) and in the 10–80 keV energy band (middle panels) are shown for MCG +8-11-11 (left-hand panel) and for NGC 6814 (right-hand panel). The ratio between 3–10 and 10–80 keV *NuSTAR* light curves is shown in the bottom panels for both sources; the red solid and dashed lines indicate the mean and standard deviation, respectively.

additional intrinsic absorbers at the redshift of the sources, which in both cases resulted to be negligible in the *NuSTAR* band. The *NuSTAR* FPMA calibration constants are fixed to 1.0 while we left the *NuSTAR* FPMB cross-calibration constants free to vary. The values found for the constant for MGC +8-11-11 and NGC 6814 are, respectively, 1.034 ± 0.009 and 0.988 ± 0.008 . These values are consistent with the expectation (Madsen et al. 2015).

3.1 X-ray/optical ratio

We used *Swift* data to compute the optical-to-X-ray spectral index (α_{ox}), defined as

$$\alpha_{\text{ox}} = -\frac{\log [L_{2\text{keV}}/L_{2500\text{\AA}}]}{2.607}. \quad (1)$$

The α_{ox} index is the slope of a hypothetical power law between 2500 Å and 2 keV rest-frame frequencies. The optical to X-ray ratio provides information about the balance between the accretion disc and the corona. The α_{ox} is found to be strongly anticorrelated with the ultraviolet luminosity density per unit frequency (see Lusso et al. 2010; Vagnetti, Antonucci & Trevese 2013 and references therein). The observed $\alpha_{\text{ox}}-L_{2500\text{\AA}}$ correlation implies that AGNs redistribute their energy in the UV and X-ray bands depending on the overall luminosity; more optical luminous AGNs emit less X-ray per unit UV luminosity than less luminous AGNs (Strateva et al. 2005).

The *Swift*/UVOT observations were analysed taking advantage of the online tool multimission archive at the ASI Science Data Center (ASDC) web site.¹ Using this tool, we performed an online interactive analysis for all the available observations for the sources MCG+8-11-11 and NGC 6814, 3 and 4, respectively. This online tool runs the standard UVOT pipeline and generates a sky map of the observation in the selected available filter. When the source is detected, it is possible to select it and run the UVOT aperture photometry. Using this tool, it is possible to compute the monochromatic flux in the selected filter. For all the observations, we extract a circular region for the source with a radius of 5 arcsec and a properly selected annular region for the background. We obtain monochromatic and extinction-corrected fluxes for MCG+8-11-11 at 3465 Å

(U), 2246 Å (UVM2), 1928 Å (UVW2) and for NGC 6814 for the same filters and UVW1 (2600 Å). Flux at 2500 Å was then computed interpolating the monochromatic flux measures for both the sources. We obtained $F_{2500\text{\AA}} = 3.98 \times 10^{-26} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$ and $F_{2500\text{\AA}} = 4.36 \times 10^{-26} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$, respectively, for MCG +8-11-11 and NGC 6814.

The 2 keV fluxes are extrapolated from the *NuSTAR* data. We found $F_{2\text{keV}} = 5.25 \pm 0.05 \times 10^{-28} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$ for MCG +8-11-11 and $F_{2\text{keV}} = 3.42 \pm 0.05 \times 10^{-29} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$ for NGC 6814.

With these values, we computed the α_{ox} using the equation (1) obtaining 1.11 for MCG +8-11-11 and 1.36 for NGC 6814.

We use the values of $L_{2500\text{\AA}}$, obtained from the previous values of the fluxes, to compute the X-ray/optical ratio with the relation found by Lusso et al. (2010)

$$\alpha_{\text{ox}}(L_{2500\text{\AA}}) = (0.154 \pm 0.0010) \log L_{2500\text{\AA}} - (3.176 \pm 0.233) \quad (2)$$

to check that the analysed sources follow the trend of the sources analysed by Lusso et al. (2010). We found $\alpha_{\text{ox}} = 1.23 \pm 0.32$ and $\alpha_{\text{ox}} = 1.51 \pm 0.25$, respectively, for MCG +8-11-11 and NGC 6814. These values are consistent, with the values of α_{ox} we found before. Moreover, the sources also follow the trend of the $L_X - L_{\text{UV}}$ of Lusso & Risaliti (2017).

3.2 MCG +8-11-11

We started the spectral analysis by fitting the 3–80 keV *NuSTAR* spectrum with a phenomenological baseline model composed of a power law with an exponential cut-off for the primary continuum (CUTOFFPL), a Gaussian line to reproduce the narrow Fe K α emission line at 6.4 keV and a cold reflection component (PEXRAV; Magdziarz & Zdziarski 1995). We fixed all element abundances to solar values and the inclination angle to $\cos i = 0.86$, $i (\sim 30^\circ)$. Since we know from the literature the presence of an emission line from H-like Fe K α (Bianchi et al. 2010), we added another narrow Gaussian line, with the centroid energy fixed to 6.966 keV. We found a $\chi^2 = 467$ for 427 degrees of freedom (d.o.f.). The residuals around the narrow Fe K α line suggest the presence of a broad-line component. We therefore added a broad Gaussian line (hereafter model A1) to the previous model. The fit slightly improved, $\chi^2/\text{d.o.f} = 446/424$. We found a resolved Fe K α line with $\sigma = 0.31^{+0.15}_{-0.20}$ keV and the centroid energy at $6.21^{+0.18}_{-0.28}$ keV. The

¹ <http://www.asdc.asi.it>

Table 1. Fit parameters for the emission features of MCG +8-11-11. Errors are at 90 per cent confidence levels.

Parameter	Fe K α NL	Fe K α BL	Fe XXVI
E (keV)	$6.40^{+0.03}_{-0.06}$	$6.21^{+0.18}_{-0.28}$	6.966^a
σ (keV)	0.0^a	$0.31^{+0.15}_{-0.20}$	0.0^a
EW (eV)	40 ± 15	90 ± 7	28^{+3}_{-7}
Flux (10^{-5} ph cm $^{-2}$ s $^{-1}$)	$2.56^{+2.5}_{-3.7}$	5.78 ± 0.9	1.4 ± 0.3

Note. ^aFixed parameter.

fitting parameters for the three lines are shown in Table 1. We found a 3–80 keV flux of $1.45 \pm 0.03 \times 10^{-10}$ erg cm $^{-2}$ s $^{-1}$. Regarding the continuum, we found a power-law index of 1.77 ± 0.02 and a cut-off energy of 260^{+190}_{-80} keV. The reflection fraction R was 0.15 ± 0.06 (Table 2, model A1).

The following step of the analysis was to replace the PEXRAV model plus the broad line with the RELXILL model (García et al. 2014), in order to test for the presence of a relativistic component (hereafter model B1). The inclination angle was fixed to a value of 30° and we fixed the ionization parameter to $\log(\frac{\xi_i}{\text{erg cm s}^{-1}}) = 0.0$ to test reflection from neutral material. Leaving the iron abundance free to vary, we found a value of iron abundance of $A_{\text{Fe}} = 3.1^{+1.7}_{-1.4}$, a reflection fraction $R = 0.24^{+0.12}_{-0.07}$ and a lower limit on the spin of the central black hole of $a > 0.6$. The $\chi^2/\text{d.o.f.}$ for this fit was 452/425. Other fitting parameters are shown in Table 2.

Leaving the ionization parameter free to vary the fit did not improve, the $\chi^2/\text{d.o.f.}$ remained the same, and we found an upper limit value for the ionization parameter of $\log(\frac{\xi_i}{\text{erg cm s}^{-1}}) < 0.05$.

We then substituted the Gaussian narrow line in the B1 model with the XILLVER model (García & Kallman 2010; García et al. 2013, hereafter model C1, see right-hand panel of Fig. 2), to test if the reprocessed spectrum could originate in distant material, like a Compton-thick torus. The photon indices the cut-off energy of the XILLVER and RELXILL models were tied together. The reflection fraction of XILLVER was free to vary. The fit gave a $\chi^2/\text{d.o.f.} = 457/425$. We found a value for the high-energy cut-off of 175^{+110}_{-50} keV (see Table 2). With this model, the observed reflection component appeared to be mostly associated with the accretion disc and to the broad part of the Fe K α . The narrow part of the Fe K α line had an emission that was relatively strong compared to the Compton hump, and the reflection fraction of the XILLVER component was found to be quite low: $R^{\text{XILLVER}} = 0.25 \pm 0.12$. This may be due to emission from a material with high iron abundance, and indeed we found an

iron overabundance in the XILLVER model of $A_{\text{Fe}} > 8.5$. The other fit parameters are reported in Table 2. Contour plots are shown in Fig. 4.

If we tried to keep tied the iron abundances between the XILLVER and the RELXILL model, we found a value of 3.2 ± 0.7 but the fit worsened significantly with a resulting $\chi^2/\text{d.o.f.} = 514/425$.

We tested if the large ratio between the line flux and the Compton hump may be due to reflection from a distant material with $N_{\text{H}} < 10^{24}$ cm $^{-2}$ using the MYTORUS model (Murphy & Yaqoob 2009). This is a more physical model with respect to the phenomenological model C1, which we used to have a measure of the cut-off. We used the MYTORUS model to fit the cold reflection and the Fe K α and Fe K β emission lines (and the associated Compton shoulders) adding the MYTORUS Compton-scattered continuum and fluorescent line tables to the RELXILL component (hereafter model D1). In the fit, we kept tied the photon indices of the MYTORUS table to the RELXILL one. The iron abundance for the RELXILL model was fixed to the values of the model C1, while for MYTORUS it is not a variable parameter being the solar value. In the standard ‘coupled’ configuration, the inclination angle for the torus was fixed to $\theta = 30^\circ$. We obtained a $\chi^2/\text{d.o.f.} = 445/435$. The best-fitting parameters are given in Table 2. We found a column density of $N_{\text{H}} = 5.0 \pm 3.0 \times 10^{23}$ cm $^{-2}$, and only a lower limit to the cut-off energy of $E_c > 250$ keV. It must be remarked that the model is not fully self-consistent in this respect because in MYTORUS the illuminating continuum must be a straight power law, with a sharp terminal energy of ~ 500 keV. However, the lower limit of the cut-off energy is consistent with the values found with the previous models. Moreover, the reflection from distant, Compton thin material is in agreement with what we found in model C1, where the relativistic reflection dominates the spectral shape above 10 keV, with respect to the non-relativistic one.

Finally, we tried to see if the two components, modelled by XILLVER and RELXILL, are not two different physical components (one arising from the accretion disc and one from a distant material: a Compton-thin gas or a gas with a supersolar abundance of iron), but two different reflection components arising from different parts of the same region, in this case the accretion disc that is in a different configuration. We modelled the spectrum with two relativistic components (two RELXILL models). Keeping all the parameters tied tighter between the two models, except for the normalization, the resulting chi-square of the fit is $\chi^2/\text{d.o.f.} = 486/426 = 1.13$. We found a lower limit on the iron abundance of $A_{\text{Fe}} > 5.88$ and there are also clear residuals around 6.4 keV which indicate the presence of

Table 2. Fitting parameters for MCG +8-11-11 for the models A1, B1, C1 and D1. Errors are at 90 per cent confidence levels. In the range of 3–0 keV, we found a flux of $1.45 \pm 0.03 \times 10^{-10}$ erg cm $^{-2}$ s $^{-1}$ and an absorption-corrected luminosity of $1.37 \pm 0.05 \times 10^{44}$ erg s $^{-1}$.

Model	PEXRAV (A1)	RELXILL (B1)	RELXILL+XILLVER (C1)	RELXILL+MYTORUS(D1)
Γ	1.77 ± 0.02	1.77 ± 0.03	1.77 ± 0.04	1.83 ± 0.03
E_c (keV)	260^{+190}_{-80}	224^{+140}_{-70}	175^{+110}_{-50}	>250
R^{PEXRAV}	0.15 ± 0.06	–	–	–
R^{RELXILL}	–	$0.24^{+0.12}_{-0.07}$	0.25 ± 0.12	0.23 ± 0.08
R^{XILLVER}	–	–	0.17 ± 0.04	–
a	–	>0.6	>0.5	<0.72
$A_{\text{Fe}}^{\text{RELXILL}}$ (Solar units)	–	$3.1^{+1.7}_{-1.4}$	<2.5	2.5^a
$A_{\text{Fe}}^{\text{XILLVER}}$ (Solar units)	–	–	>8.5	–
$N_{\text{H}}^{\text{MYT}}$ (10^{23} cm $^{-2}$)	–	–	–	5.0 ± 3.0
$\chi^2/\text{d.o.f.}$	1.05	1.06	1.07	1.05

Note. ^aFixed parameter.

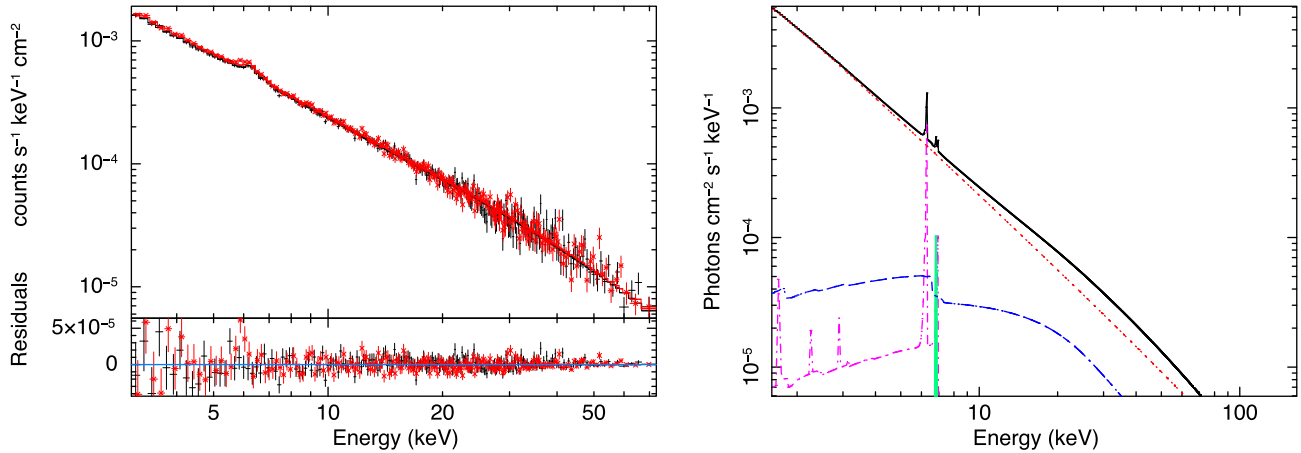


Figure 2. Left-hand panel: data and best-fitting model of MCG +8-11-11 extrapolated from *NuSTAR* FPMA (black dot) and FPMB (red asterisk) spectra when model C1 (right-hand panel) was used. Right-hand panel: best-fitting model (black solid line); the blue dashed line is the RELXILL component, the red dotted line represents the primary continuum, the bold solid green line represents the Fe XXVI emission line and the magenta dash-dotted line is the XILLVER component. See the text for more details.

the narrow component of the Fe K α line. Adding this component and allowing to vary the iron abundance parameter, the reflection fraction parameter and the black hole spin parameter, we found a fit that is statistically equivalent to our best-fitting model, with a chi-square of $\chi^2/\text{d.o.f.} = 452/424 = 1.07$. We found one RELXILL component with an upper limit on the iron abundance of $A_{\text{Fe}} < 1.48$, an upper limit on the spin parameter of $a < 0.48$ and a reflection fraction $R = 0.16 \pm 0.05$. The second RELXILL component showed a lower limit on the iron abundance of $A_{\text{Fe}} > 6.35$, a spin value of $a = 0.07 \pm 0.01$ and a reflection fraction $R < 0.07$. This suggested us that the latter is not a relativistic component and it could not be produced by a Compton thick material since it had a very low associated reflection component and high iron abundance, similarly to the XILLVER component of our best fit (model C1). This scenario is consistent with what Bianchi et al. (2010) and Mantovani, Nandra & Ponti (2016) found for this source.

Fluxes and centroid energies of the broad and narrow Fe K α lines and of the Fe XXVI line resulted to be the same, within the errors, among the various models previously described.

3.2.1 Coronal parameters

Assuming that the primary emission is due to Comptonization of thermal disc photons in a hot corona, we estimated the coronal parameters substituting the cut-off power law with a Comptonization model. We modelled the relativistic reflection and the non-relativistic reflection using, respectively, RELXILLCP and XILLVER-COMP models (García & Kallman 2010; Dauser et al. 2014; García et al. 2014). These models use the NTHCOMP model (Zdziarski, Johnson & Magdziarz 1996; Życki, Done & Smith 1999) for the incident continuum. In RELXILLCP and XILLVER-COMP models, the maximum temperature of disc blackbody photons (which serve as seeds for Comptonization) is 0.1 keV, and it is not allowed to vary. The fit gave a $\chi^2/\text{d.o.f.} = 465/426 = 1.09$. We found a photon index value of $1.84_{-0.05}^{+0.03}$; we also found a lower limit for the iron abundance of the non-relativistic reflection component: $A_{\text{Fe}} > 8.01$ solar unit, and an upper limit for the iron abundance of the relativistic reflection component: $A_{\text{Fe}} < 0.5$ solar unit, in agreement with the values reported in Table 2. We found a coronal temperature of 60_{-30}^{+110} keV,

roughly in agreement with the expected relation $E_c = 2-3 \times kT_e$ (Petrucci et al. 2000, 2001).

3.3 NGC 6814

The fitting procedure for NGC 6814 was very similar to that adopted for MCG +8-11-11. We started the analysis by fitting the 3–80 keV *NuSTAR* spectrum with the phenomenological baseline model, previously described in Section 3.2 and composed by a cut-off power law, a PEXRAV plus a narrow Gaussian line around 6.4 keV. Also in this case, we fixed all element abundances to solar values and the inclination angle to $\cos i = 0.86$ ($i \sim 30^\circ$). Since Walton et al. (2013) found the presence of an emission line from H-like Fe K α with an EW of 90_{-40}^{+30} eV, we added also a narrow Gaussian line with the centroid energy fixed to 6.966 keV. However, only an upper limit to the EW of < 15 eV was found, which was more than four times lower than the value found by Walton et al. (2013). We also found an upper limit for the line flux of $< 2.1 \times 10^{-6}$ ph cm $^{-2}$. This was probably due to the fact that the line was diluted by a continuum that was about four times higher than that found by (Walton et al. 2013, see below). We therefore did not include this line in this and in the following fits. We found a $\chi^2/\text{d.o.f.} = 398/348$. The residuals around the narrow Fe K α line suggested the presence of a broad line. Adding a broad Gaussian line at the PEXRAV plus the narrow Fe K α line model (model A2) around 6.4 keV, we found a resolved Fe K α line with $\sigma = 0.59_{-0.21}^{+0.37}$ keV, an EW of 102 ± 15 eV and flux of $4.2 \pm 1.4 \times 10^{-5}$ ph cm $^{-2}$ s $^{-1}$. The centroid energy of the Fe K α line was found to be at $6.43_{-0.06}^{+0.03}$ keV. The line had an EW of 70 ± 7 eV, consistent with what Walton et al. (2013) found. The flux of the line was $2.7 \pm 0.7 \times 10^{-5}$ ph cm $^{-2}$ s $^{-1}$. The chi square for the fit with the model A2 was $\chi^2/\text{d.o.f.} = 373/345$. The high-energy cut-off was 260_{-80}^{+220} keV and the reflection fraction $R = 0.15 \pm 0.07$. The 3–80 keV flux was $1.04 \pm 0.04 \times 10^{-10}$ erg cm $^{-2}$ s $^{-1}$. Other parameters are reported in Table 3.

The following step in the analysis was to replace the PEXRAV model plus the broad line with the RELXILL model, as for the previous source but without the Fe XXVI line (hereafter model B2) in order to test for the presence of a relativistic component. The inclination angle was fixed to 30° and the ionization parameter was fixed

Table 3. Fitting parameters for NGC 6814 using the models A2, B2, C2 and D2, as described in the text. Errors are at 90 per cent confidence levels. In the range of 3–0 keV, we found a flux of $1.04 \pm 0.04 \times 10^{-10}$ erg cm $^{-2}$ s $^{-1}$ and an absorption-corrected luminosity of $6.21 \pm 0.12 \times 10^{42}$ erg s $^{-1}$.

Model	PEXRAV (A2)	RELXILL (B2)	RELXILL+XILLVER (C2)	RELXILL+MYTORUS(D2)
Γ	1.72 ± 0.02	1.74 ± 0.02	$1.71^{+0.04}_{-0.03}$	1.80 ± 0.02
E_c (keV)	260^{+220}_{-80}	210^{+80}_{-50}	155^{+70}_{-35}	>260
R^{PEXRAV}	0.15 ± 0.07	–	–	–
R^{RELXILL}	–	0.26 ± 0.1	$0.27^{+0.10}_{-0.12}$	$0.26^{+0.18}_{-0.11}$
R^{XILLVER}	–	–	0.17 ± 0.03	–
a	–	>0.2	>0.03	>0.4
$A_{\text{Fe}}^{\text{RELXILL}}$ (solar units)	–	<1.4	$1.8^{+0.8}_{-1.3}$	1.8^a
$A_{\text{Fe}}^{\text{XILLVER}}$ (solar units)	–	–	>7.00	–
$N_{\text{H}}^{\text{MYT}}$ (10^{23} cm $^{-2}$)	–	–	–	$3.5^{+3.0}_{-2.0}$
$\frac{\chi^2}{\text{d.o.f}}$	1.08	1.06	1.08	1.06

Note. ^aFixed parameter.

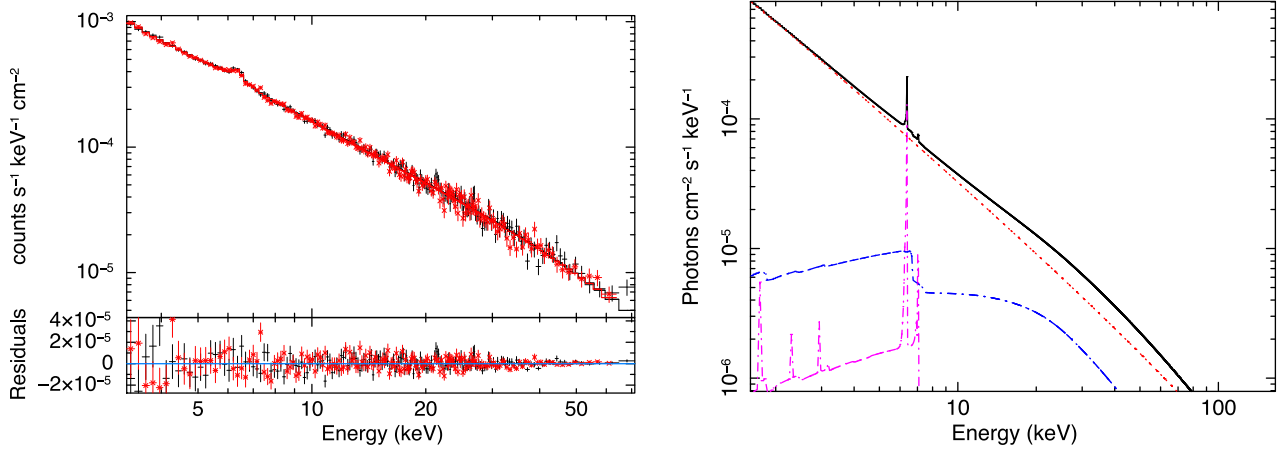


Figure 3. Left-hand panel: data and best-fitting model of NGC 6814 extrapolated from *NuSTAR* FPMA (black dot) and FPMB (red asterisk) spectra when model C2 (right-hand panel) is used. Right-hand panel: best-fitting model (solid black line); the blue dashed line is the RELXILL component, the red dotted line represents the primary continuum and the magenta dash-dotted line is the XILLVER component. See the text for more details.

to $\log(\frac{\xi_i}{\text{erg cm s}^{-1}}) = 0.0$. Leaving the iron abundance free to vary, we found a value of $A_{\text{Fe}} < 1.4$ in solar units, a reflection fraction $R = 0.26 \pm 0.1$ and a lower limit on the spin of the central black hole of $a > 0.2$. The χ^2 values in this case was $\chi^2/\text{d.o.f.} = 366/343$. Leaving the ionization parameter free to vary, the fit did not improve, the χ^2 remained the same and we found an upper limit value for the ionization parameter of $\log(\frac{\xi_i}{\text{erg cm s}^{-1}}) < 0.34$. Other fitting parameters are shown in Table 3.

We then replaced the Gaussian narrow line in the model B2 with the XILLVER model (hereafter model C2; see right-hand panel of Fig. 3) to test if the reprocessed spectrum could originate in a distant material, as it was done for the other source in the previous section. The photon indices the cut-off energy of the XILLVER and RELXILL models are tied together. The reflection fraction of XILLVER was free to vary. The fit gave a $\chi^2/\text{d.o.f.} = 375/346$. We found a value for the high-energy cut-off of 155^{+70}_{-35} keV (see Table 3). The other fit parameters are reported in Table 3. Contour plots are shown in Fig. 4. Also for NGC 6814, as for MCG +8-11-11, we found an iron overabundance in the XILLVER model, $A_{\text{Fe}} > 7.0$ due to an emission of the narrow part of the Fe K α line which was relatively strong compared to the Compton hump. If we tried to keep tied the iron abundances between the XILLVER and the RELXILL model, we found a lower limit value of $A_{\text{Fe}} > 4$

and the fit worsened significantly with a resulting $\chi^2/\text{d.o.f.} = 379/347$.

Similarly to the case of MCG +8-11-11, we tested the alternative model in which XILLVER was replaced by reflection from distant material with $N_{\text{H}} < 10^{24}$ cm $^{-2}$ which could reproduce the narrow part of the fluorescence emission line from the iron K-shell with a small Compton hump. We used the MYTORUS model as described in Section 3.2 (hereafter model D2). We obtained a $\chi^2/\text{d.o.f.} = 368/347$. The best-fitting parameters are given in Table 3. We found a column density of $N_{\text{H}} = 3.5^{+3.0}_{-2.0} \times 10^{23}$ cm $^{-2}$. Again, we found only a lower limit to the cut-off energy, $E_c > 260$ keV, which was consistent with the values found with the previous models.

Yamauchi et al. (1992) found an iron overabundance in the *Ginga* spectrum of the source. They justified the fact with the presence of a partially ionized state that may give an ‘apparent’ overabundance because the partially ionized gas is transparent for the soft X-rays but absorbs X-rays above the Fe K-edge energy. In our fit, we considered neutral reflection material, so the supersolar value for the iron abundance could not be an effect of the ionization. Even in this case, we tested if the two reflection components could arise from two different parts of the accretion disc if it is in a different configuration. We modelled the spectrum with two RELXILL models. First, we kept all the parameters tied tighter between the two models,

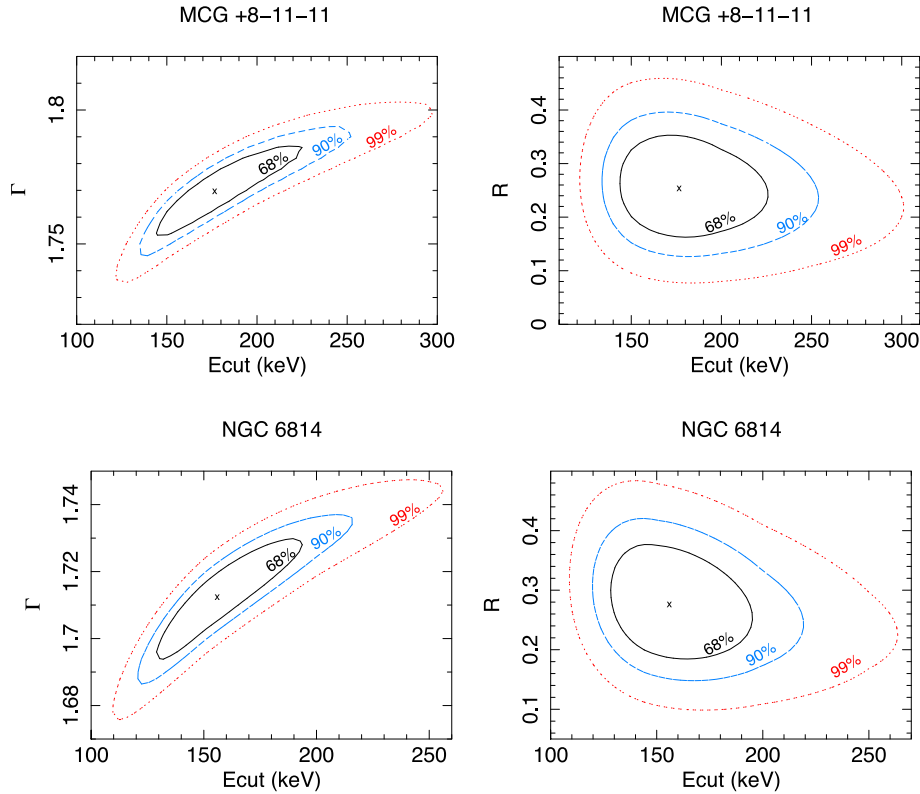


Figure 4. E_c - Γ contour plot (left-hand panels) and E_c - R contour plot (right-hand panels) for MCG +8-11-11 (top panels) and NGC 6814 (lower panels). The black solid, blue dashed and red dotted curves refer to the 68, 90 and 99 per cent confidence levels, respectively. The X represents the best-fitting value of the parameters.

except for the normalization, the resulting chi-square of the fit is $\chi^2/\text{d.o.f.} = 409/347 = 1.17$. We found a lower limit on the iron abundance of $A_{\text{Fe}} > 5.88$ and there are also clear residuals around 6.4 keV which indicate the presence of the narrow component of the Fe K α line. Adding this component and allowing to vary the iron abundance parameter, the reflection fraction parameter and the black hole spin parameter, we found a fit with a chi-square of $\chi^2/\text{d.o.f.} = 373/343 = 1.09$. We found one RELXILL component with an upper limit on the iron abundance of $A_{\text{Fe}} < 2.09$, an upper limit on the spin parameter of $a < 0.48$ and a reflection fraction $R = 0.27 \pm 0.12$. The second RELXILL component showed a lower limit on the iron abundance of $A_{\text{Fe}} > 6.35$, an upper limit on the spin value $a < 0.15$ and a lower limit on the reflection fraction $R < 0.12$. Also, NGC 6814 shows one of the RELXILL components that is not a relativistic component and it could not be produced by a Compton thick material since it had a very low associated reflection component and high iron abundance, similarly to the XILLVER component of our best fit (model C2). Fluxes of the broad and narrow Fe K α lines and of the Fe XXVI line were the same, within the errors, among the various models previously described.

3.3.1 Coronal parameters

The final step was to use Comptonization models to estimate the coronal parameters as in Section 3.2.1. We used RELXILLCP and XILLVER-COMP to model both the relativistic and non-relativistic reflection spectrum, see Section 3.2.1. The fit gave a $\chi^2/\text{d.o.f.} = 387/347 = 1.11$. We found a photon index value of 1.79 ± 0.03 and a coronal temperature of 45_{-17}^{+100} keV, again in agreement with the $E_c = 2-3 \times kT_c$ relation (Petrucci et al. 2000, 2001). We also found a lower limit for the iron abundance of the non-relativistic reflection

component: $A_{\text{Fe}} > 7.89$ solar unit, and an upper limit for the iron abundance of the relativistic reflection component: $A_{\text{Fe}} < 0.73$ solar unit, in agreement with the values reported in Table 3.

4 DISCUSSION AND CONCLUSIONS

We have presented the analysis of the *NuSTAR* observations of the Seyfert 1 galaxies NGC 6814 and MCG +8-11-11.

The 2–10 keV absorption-corrected luminosities from the *NuSTAR* observations of the two sources are $L_{2-10} = 2.04 \times 10^{42}$ erg s $^{-1}$ for NGC 6814 and $L_{2-10} = 5.13 \times 10^{43}$ erg s $^{-1}$ for MCG +8-11-11. Using the 2–10 keV bolometric correction of Marconi et al. (2004), we estimated the bolometric luminosity to be $L_{\text{bol}} = 0.24 \times 10^{44}$ erg s $^{-1}$ (NGC 6814) and $L_{\text{bol}} = 14.2 \times 10^{44}$ erg s $^{-1}$ (MCG +8-11-11). From this bolometric luminosity, with the black hole masses of $\log \frac{M_{\text{BH}}}{M_{\odot}} = 7.19 \pm 0.02$ (Bian & Zhao 2003, MCG +8-11-11) and $\log \frac{M_{\text{BH}}}{M_{\odot}} = 6.99_{-0.25}^{+0.32}$ (Pancoast et al. 2014, 2015, NGC 6814), we estimated the Eddington ratio to be 2.46×10^{-3} for NGC 6814 and 7.54×10^{-1} for MCG +8-11-11.

Thanks to the *NuSTAR* sensitivity at high energies, it was possible to measure the high-energy cut-off value for both sources; we found 175_{-50}^{+110} and 155_{-35}^{+70} keV, respectively, for MCG +8-11-11 and NGC 6814. We also found a disc reflection component and we constrained the reflection fraction finding 0.25 ± 0.12 for MCG +8-11-11 and of $0.27_{-0.12}^{+0.10}$ for NGC 6814.

Both sources showed a slightly broadened relativistic Fe K α line plus a narrow component. The reflection component was modest, and mostly associated with the broad-line component. The low reflection fraction found in MCG +8-11-11 was consistent with

the value found by Mantovani et al. (2016). Past observations of NGC 6814 did not show a broad line (Bentz et al. 2009; Malizia et al. 2014; Ricci et al. 2014). The Compton hump associated with the narrow line was very small, similarly to what found in another Seyfert galaxy, NGC 7213 (Ursini et al. 2015).

We found in our analysis a slightly broadened relativistic Fe K α line plus a narrow component. The former is modelled by the relativistic model RELXILL while the latter by the non-relativistic model XILLVER. The iron abundance measured with the RELXILL model is considerably lower with respect to the supersolar abundance measured with the XILLVER model. Since there was no Compton reflection hump associated with the narrow component of the iron K α line, it could not be produced by a Compton-thick material, like the accretion disc or the Compton-thick torus, and thus almost all the reflection should be associated with the accretion disc.

The interaction of X-rays with a material with supersolar abundance of iron gives rise to a reflection component with a small Compton hump associated with the narrow iron K α emission line. This is because the iron atoms interact with the X-rays, causing photoelectric absorption and the spectrum shows an iron K α line with a drop at ~ 7 keV, due to the photoelectric absorption edge. The depth of the edge increases with the iron abundance and saturates around $A_{\text{Fe}} \sim 10$ (Matt et al. 1997). The curvature of the continuum above ~ 10 becomes weaker, resulting in a very small Compton reflection.

A reflection component with a small Compton hump associated with the narrow iron K α emission line could be produced also by the interaction of X-rays with a Compton thin material. To confirm this scenario, we fitted the cold reflection component with the MYTORUS model and we found values of $N_{\text{H}} < 10^{24}$ cm $^{-2}$. This could imply that there would be no evidence of the classical Compton-thick torus in these sources, as already suggested by Bianchi et al. (2010) for MCG +8-11-11.

Ultimately, we also tested if the two different reflection components should be part of the emission by the same material, such as an accretion disc that is in a different configuration. But this attempt has further strengthened the scenario described above.

Regarding the relativistic component, for both sources we derived only lower or upper limits to the spin of the black hole of the two sources, depending on the adopted model. The relativistic line was fairly broad, so there are degeneracies in the parameters of the models which prevent us from determining a good measurement of the spin. We conclude that this parameter is basically unconstrained, not surprisingly given the relative weakness of the relativistic reflection.

We estimated the coronal parameters by fitting the NuSTAR data with a model that takes into account both the relativistic and non-relativistic reflection when illuminated by a thermally Comptonized continuum. We found a coronal temperature of 60_{-30}^{+110} keV for MCG +8-11-11 and 45_{-17}^{+100} keV for NGC 6814. It is interesting to note that the coronal temperatures of the two sources are very similar despite an order of magnitude difference in mass and Eddington ratio.

We estimated the optical depth using the relation from Beloborodov (1999):

$$\Gamma \approx \frac{9}{4} y^{-2/9}, \quad (3)$$

where Γ is the photon index of the spectrum between 2 and 10 keV. The dependence from the optical depth is in the relativistic y -parameter:

$$y = 4 (\Theta_e + 4\Theta_e^2) \tau (\tau + 1), \quad (4)$$

Table 4. Coronal parameters for MCG +8-11-11 and NGC 6814 when the self-consistent model XILLVER-COMP + RELXILLCP is used to fit the data. The optical depths are extrapolated from Beloborodov (1999). Errors are at 90 percent confidence levels.

Parameter	MCG +8-11-11	NGC 6814
kT_e (keV)	60_{-30}^{+110}	45_{-17}^{+100}
τ	1.8 ± 0.2	2.5 ± 0.2
Γ	$1.84_{-0.05}^{+0.03}$	1.79 ± 0.03
$A_{\text{Fe}}^{\text{RELXILL}}$ (solar units)	< 0.5	< 0.73
$A_{\text{Fe}}^{\text{XILLVER}}$ (solar units)	> 8.01	> 7.89
$\frac{\chi^2}{\text{d.o.f}}$	1.09	1.11

where Θ_e is the electron temperature normalized to the electron rest energy:

$$\Theta_e = \frac{kT_e}{m_e c^2}. \quad (5)$$

We found $\tau = 1.8 \pm 0.2$ for MCG +8-11-11 and $\tau = 2.5 \pm 0.2$ for NGC 6814.

We used the values of the coronal temperature reported in Table 4 to put MCG +8-11-11 and NGC 6814 in the compactness–temperature (Θ_e – ℓ) diagram (Fabian et al. 2015, and references therein). Here, Θ_e is the electron temperature normalized to the electron rest energy, defined above in equation (5), and ℓ is the dimensionless compactness parameter (Fabian et al. 2015):

$$\ell = \frac{L}{R} \frac{\sigma_{\text{T}}}{m_e c^3}, \quad (6)$$

where L is the luminosity and R is the radius of the corona (assumed spherical). To compute the compactness parameter, following Fabian et al. (2015) we adopted the luminosity of the power-law component extrapolated to the 0.1–200 keV band; since no measurement exists for the corona radius, we assume a value of 10 gravitational radii R_g . For MCG +8-11-11, we found $\ell = 27 \pm 12 (R_{10})^{-1}$ and $\Theta_e = 0.11_{-0.10}^{+0.15}$; for NGC 6814 $\ell = 14.5 \pm 4.5 (R_{10})^{-1}$ and $\Theta_e = 0.08_{-0.05}^{+0.1}$. Here, R_{10} is the ratio between the corona radius and $10 R_g$.

With these values of Θ_e and ℓ , both sources are positioned under the Svensson pair runaway line for a spherical geometry (Svensson 1984), above the $e^- - e^-$ coupling line, like most of the sources among those analysed by Fabian et al. (2017). The pair runaway line in the Θ_e – ℓ diagram is a curve that determines a forbidden region in which the pair production exceeds the annihilation. The detailed shape of this line depends on the source of soft photons and on the radiation mechanism. Above the $e^- - e^-$ coupling line, the $e^- - e^-$ coupling time-scale is longer than the Compton cooling time-scale. The location of these sources within the Θ_e – ℓ plane fits well in the scenario in which the AGN spectral shape is controlled by pair production and annihilation.

ACKNOWLEDGEMENTS

We thank the anonymous referee for comments that helped improving the clarity of the paper. This work made use of data from the NuSTAR mission, a project led by the California Institute of Technology, managed by the Jet Propulsion Laboratory and funded by the National Aeronautics and Space Administration. We thank the NuSTAR Operations, Software and Calibration teams for support with the execution and analysis of these observations. This

research has made use of the NuSTARDAS jointly developed by the ASDC (Italy) and the California Institute of Technology (USA). AT, AM and GM acknowledge financial support from Italian Space Agency under grant ASI/INAF I/037/12/0-011/13, and SB under grant ASI-INAF I/037/12/P1. Part of this work is based on archival data, software or online services provided by the ASDC. AT, SB, AM, GM and RM acknowledge financial support from the European Union Seventh Framework Programme (FP7/2007-2013) under grant agreement no. 312789. POP acknowledges financial support from the CNES and the CNRS/PNHE. EP and SB acknowledge financial contribution from the agreement ASI-INAF I/037/12/0.

REFERENCES

- Arnaud K. A., 1996, in Jacoby G. H., Barnes J., eds, ASP Conf. Ser. Vol. 101, *Astronomical Data Analysis Software and Systems V*. Astron. Soc. Pac., San Francisco, p. 17
- Beloborodov A. M., 1999, in Poutanen J., Svensson R., eds, ASP Conf. Ser. Vol. 161, *High Energy Processes in Accreting Black Holes*. Astron. Soc. Pac., San Francisco, p. 295
- Bentz M. C., Peterson B. M., Netzer H., Pogge R. W., Vestergaard M., 2009, *ApJ*, 697, 160
- Bian W., Zhao Y., 2003, *MNRAS*, 343, 164
- Bianchi S., de Angelis I., Matt G., La Parola V., de Rosa A., Grandi P., Jiménez B. E., Piconcelli E., 2010, *A&A*, 522, A64
- Blackburn J. K., 1995, in Shaw R. A., Payne H. E., Hayes J. J. E., eds, ASP Conf. Ser. Vol. 77, *Astronomical Data Analysis Software and Systems IV*. Astron. Soc. Pac., San Francisco, p. 367
- Dauser T., García J., Parker M. L., Fabian A. C., Wilms J., 2014, *MNRAS*, 444, L100
- Fabian A. C., Rees M. J., Stella L., White N. E., 1989, *MNRAS*, 238, 729
- Fabian A. C. et al., 2002, *MNRAS*, 335, L1
- Fabian A. C. et al., 2009, *Nature*, 459, 540
- Fabian A. C., Lohfink A., Kara E., Parker M. L., Vasudevan R., Reynolds C. S., 2015, *MNRAS*, 451, 4375
- Fabian A. C., Lohfink A., Belmont R., Malzac J., Coppi P., 2017, *MNRAS*, 467, 2566
- García J., Kallman T. R., 2010, *ApJ*, 718, 695
- García J., Dauser T., Reynolds C. S., Kallman T. R., McClintock J. E., Wilms J., Eikmann W., 2013, *ApJ*, 768, 146
- García J. et al., 2014, *ApJ*, 782, 76
- George I. M., Fabian A. C., 1991, *MNRAS*, 249, 352
- Grandi P., Haardt F., Ghisellini G., Grove E. J., Maraschi L., Urry C. M., 1998, *ApJ*, 498, 220
- Haardt F., Maraschi L., 1991, *ApJ*, 380, L51
- Haardt F., Maraschi L., 1993, *ApJ*, 413, 507
- Haardt F., Matt G., 1993, *MNRAS*, 261, 346
- Haardt F., Maraschi L., Ghisellini G., 1994, *ApJ*, 432, L95
- Harrison F. A. et al., 2013, *ApJ*, 770, 103
- Kalberla P. M. W., Burton W. B., Hartmann D., Arnal E. M., Bajaja E., Morras R., Pöppel W. G. L., 2005, *A&A*, 440, 775
- Lusso E., Risaliti G., 2017, *A&A*, 602, A79
- Lusso E. et al., 2010, *A&A*, 512, A34
- Madsen K. K. et al., 2015, *ApJS*, 220, 8
- Magdziarz P., Zdziarski A. A., 1995, *MNRAS*, 273, 837
- Malizia A., Bassani L., Bazzano A., Bird A. J., Masetti N., Panessa F., Stephen J. B., Ubertini P., 2012, *MNRAS*, 426, 1750
- Malizia A., Molina M., Bassani L., Stephen J. B., Bazzano A., Ubertini P., Bird A. J., 2014, *ApJ*, 782, L25
- Mantovani G., Nandra K., Ponti G., 2016, *MNRAS*, 458, 4198
- Marconi A., Risaliti G., Gilli R., Hunt L. K., Maiolino R., Salvati M., 2004, *MNRAS*, 351, 169
- Marinucci A., Tortosa A., NuSTAR AGN Physics Working Group, 2016, *Astron. Nachr.*, 337, 490
- Matt G., Perola G. C., Piro L., 1991, *A&A*, 247, 25
- Matt G., Fabian A. C., Reynolds C. S., 1997, *MNRAS*, 289, 175
- Matt G., Bianchi S., de Rosa A., Grandi P., Perola G. C., 2006, *A&A*, 445, 451
- Molina M. et al., 2009, *MNRAS*, 399, 1293
- Mukai K., Hellier C., Madejski G., Patterson J., Skillman D. R., 2003, *ApJ*, 597, 479
- Murphy K. D., Yaqoob T., 2009, *MNRAS*, 397, 1549
- Nandra K., Pounds K. A., 1994, *MNRAS*, 268, 405
- Pancoast A., Brewer B. J., Treu T., Park D., Barth A. J., Bentz M. C., Woo J.-H., 2014, *MNRAS*, 445, 3073
- Pancoast A., Brewer B. J., Treu T., Park D., Barth A. J., Bentz M. C., Woo J.-H., 2015, *MNRAS*, 448, 3070
- Perola G. C. et al., 2000, *A&A*, 358, 117
- Petrucci P. O. et al., 2000, *ApJ*, 540, 131
- Petrucci P. O. et al., 2001, *ApJ*, 556, 716
- Reynolds C. S., Nowak M. A., 2003, *Phys. Rep.*, 377, 389
- Ricci C., Ueda Y., Ichikawa K., Paltani S., Boissay R., Gandhi P., Stalewski M., Awaki H., 2014, *A&A*, 567, A142
- Shapiro S. L., Lightman A. P., Eardley D. M., 1976, *ApJ*, 204, 187
- Strateva I. V., Brandt W. N., Schneider D. P., Vanden Berk D. G., Vignali C., 2005, *AJ*, 130, 387
- Sunyaev R. A., Titarchuk L. G., 1980, *A&A*, 86, 121
- Svensson R., 1984, *MNRAS*, 209, 175
- Tanaka Y. et al., 1995, *Nature*, 375, 659
- Ursini F. et al., 2015, *MNRAS*, 452, 3266
- Vagnetti F., Antonucci M., Trevese D., 2013, *A&A*, 550, A71
- Walton D. J. et al., 2013, *ApJ*, 777, L23
- Yamauchi M., Matsuoka M., Kawai N., Yoshida A., 1992, *ApJ*, 395, 453
- Zdziarski A. A., Johnson W. N., Magdziarz P., 1996, *MNRAS*, 283, 193
- Życki P. T., Done C., Smith D. A., 1999, *MNRAS*, 309, 561

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.