



HAL
open science

Interstellar comet 2I/Borisov as seen by MUSE: C 2 , NH 2 and red CN detections

Michele T Bannister, Cyrielle Opitom, Alan Fitzsimmons, Youssef Moulane, Emmanuel Jehin, Darryl Seligman, Philippe Rousselot, Matthew M Knight, Michael Marsset, Megan E Schwamb, et al.

► To cite this version:

Michele T Bannister, Cyrielle Opitom, Alan Fitzsimmons, Youssef Moulane, Emmanuel Jehin, et al.. Interstellar comet 2I/Borisov as seen by MUSE: C 2 , NH 2 and red CN detections. 2020. hal-03010656

HAL Id: hal-03010656

<https://hal.science/hal-03010656v1>

Preprint submitted on 17 Nov 2020

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Interstellar comet 2I/Borisov as seen by MUSE: C₂, NH₂ and red CN detections

MICHELE T. BANNISTER,^{1,2} CYRIELLE OPITOM,^{3,4} ALAN FITZSIMMONS,¹ YOUSSEF MOULANE,^{3,5,6} EMMANUEL JEHIN,⁵
DARRYL SELIGMAN,⁷ PHILIPPE ROUSSELOT,⁸ MATTHEW M. KNIGHT,^{9,10} MICHAEL MARSETT,¹¹ MEGAN E. SCHWAMB,¹
AURÉLIE GUILBERT-LEPOUTRE,¹² LAURENT JORDA,¹³ PIERRE VERNAZZA,¹³ AND ZOUHAIR BENKHALDOUN⁶

¹*Astrophysics Research Centre, School of Mathematics and Physics, Queen's University Belfast, Belfast BT7 1NN, United Kingdom*

²*School of Physical and Chemical Sciences – Te Kura Matū, University of Canterbury, Private Bag 4800, Christchurch 8140, New Zealand*

³*ESO (European Southern Observatory) - Alonso de Cordova 3107, Vitacura, Santiago Chile*

⁴*Institute for Astronomy, University of Edinburgh, Royal Observatory, Edinburgh EH9 3HJ, UK*

⁵*STAR Institute, Université de Liège, Allée du 6 aout, 19C, 4000 Liège, Belgium*

⁶*Oukaimeden Observatory, Cadi Ayyad University, Marrakech, Morocco*

⁷*Department of Astronomy, Yale University, 52 Hillhouse Ave., New Haven, CT 06517*

⁸*Institut UTINAM UMR 6213, CNRS, Univ. Bourgogne Franche-Comt, OSU THETA, BP 1615, 25010 Besancon Cedex, France*

⁹*Department of Physics, United States Naval Academy, 572C Holloway Rd, Annapolis, MD 21402, USA*

¹⁰*University of Maryland, Department of Astronomy, College Park, MD 20742, USA*

¹¹*Department of Earth, Atmospheric and Planetary Sciences, MIT, 77 Massachusetts Avenue, Cambridge, MA 02139, USA*

¹²*Laboratoire de Géologie de Lyon, LGL-TPE, UMR 5276 CNRS / Université de Lyon / Université Claude Bernard Lyon 1 / ENS Lyon, 69622 Villeurbanne, France*

¹³*Aix Marseille Univ, CNRS, LAM, Laboratoire d'Astrophysique de Marseille, Marseille, France*

(Received 29 Jan 2020; Revised; Accepted)

Submitted to ApJ Letters

ABSTRACT

We report the clear detection of C₂ and of abundant NH₂ in the first prominently active interstellar comet, 2I/Borisov. We observed 2I on three nights in November 2019 at optical wavelengths 4800–9300 Å with the Multi-Unit Spectroscopic Explorer (MUSE) integral-field spectrograph on the ESO/Very Large Telescope. These data, together with observations close in time from both 0.6-m TRAPPIST telescopes, provide constraints on the production rates of species of gas in 2I's coma. From the MUSE detection on all epochs of several bands of the optical emission of the C₂ Swan system, a rich emission spectrum of NH₂ with many highly visible bands, and the red (1-0) bandhead of CN, together with violet CN detections by TRAPPIST, we infer production rates of $Q(\text{C}_2) = 1.1 \times 10^{24} \text{ mol s}^{-1}$, $Q(\text{NH}_2) = 4.8 \times 10^{24} \text{ mol s}^{-1}$ and $Q(\text{CN}) = (1.8 \pm 0.2) \times 10^{24} \text{ mol s}^{-1}$. In late November at 2.03 au, 2I had a production ratio of C₂/CN = 0.61, only barely carbon-chain depleted, in contrast to earlier reports measured further from the Sun of strong carbon-chain depletion. Thus, 2I has shown evolution in its C₂ production rate: a parent molecule reservoir has started sublimating. At $Q(\text{NH}_2)/Q(\text{CN}) = 2.7$, this second interstellar object is enriched in NH₂, relative to the known Solar System sample.

1. INTRODUCTION

Interstellar objects offer an unprecedented opportunity to closely study remnants of the star and planet formation process from systems other than our own. The comae of interstellar comets permit detailed characterization of the chemical composition of their volatile material and the physical properties of their dust. Although planet formation is a robust process, the architectures of extrasolar planetary systems are remarkably

diverse; many are substantially different to the architecture of our own system. Characterizing planetesimals from other stars will contextualize the formation and evolution of our own system, through comparison to the Solar System minor planet populations.

2I/Borisov¹ is the second interstellar object ever detected entering our Solar System, and the first to exhibit distinct cometary activity. In contrast to 2I, the first interstellar object, 1I/'Oumuamua, displayed a marked lack of a coma in deep stacks of images (e.g. Meech

Corresponding author: Michele Bannister
michele.t.bannister@gmail.com

¹ Initially designated C/2019 Q4 (Borisov); MPEC 2019-R106
<https://www.minorplanetcenter.net/mpec/K19/K19RA6.html>

et al. 2017). The sublimation of volatiles from 1I was only inferred from the minor non-gravitational acceleration in its outbound trajectory (Micheli et al. 2018). However, 1I had distinctive characteristics even from the upper-limit measurements of its volatiles: for instance, it needed to be depleted in CN by at least a factor of 15 relative to typical abundances in comets, if the non-gravitational forces were due to the sublimation of H₂O (‘Oumuamua ISSI Team et al. 2019). Only two previous Solar System comets have been seen with comparable CN depletions, and those were hypothesized to be potentially captured interstellar objects (Fink 1992; Schleicher 2008). 2I has a definitively interstellar origin, with an escape velocity, $v_\infty = 32.304 \pm 0.001$ km/s, on a hyperbolic trajectory with an eccentricity of 3.3575 ± 0.0003^2 . This is well within the expected distribution of values for the eccentricity of a detectable active interstellar object, given the inherent bias in the Sun’s gravitational focusing of the trajectories of these objects (Engelhardt et al. 2017). 2I entered the Solar System at an inclination of 44° to the ecliptic, and its inbound trajectory passed outside of the orbit of Mars at 2.006 au, near the peak of detectable perihelion values for active objects (Engelhardt et al. 2017). Hubble Space Telescope optical imaging indicates a nuclear size of 0.2–0.5 km for 2I (Jewitt et al. 2019), supported by upper limits from ground-based infrared imaging (Bolin et al. 2019; Lee et al. 2019). From this, the size distribution of interstellar objects is such that the discovery of 2I is consistent with extrapolations based on 1I (Jewitt et al. 2019).

The cometary activity of 2I provides the first opportunity for an extensive compositional study of a volatile-rich interstellar object. A dusty coma with an initial colour of $g' - r' = 0.63 \pm 0.03$ and D-type visible spectrum implied a first-look similarity to long-period comets (Guzik et al. 2019; de León et al. 2019; Jewitt & Luu 2019; Bolin et al. 2019; Yang et al. 2019). High-airmass spectra soon after the 2019 August 30 discovery showed the presence of CN (Fitzsimmons et al. 2019). Ongoing upper limits on C₂ (Opitom et al. 2019b; Kareta et al. 2019), edging towards weak detections (Lin et al. 2019), have been at levels implying that 2I is carbon-depleted, similar to ~30% of the Solar System comets (A’Hearn et al. 1995; Cochran et al. 2012). The brightening rate in Ye et al. (2019)’s pre-discovery observations of 2I back to December 2018 appears flat, as does the more densely sampled brightening rate since discovery (e.g. Fig. 3, Opitom et al. 2019b). This is consistent with the behaviour of dynamically new comets, and much less steep than that of Jupiter-family comets (Whipple 1978). 2I’s lack of detection in Ye et al. (2019)’s earlier imaging implies that activity driven by

species more volatile than H₂O, such as CO or CO₂, began between ~7.8–8.5 au. H₂O sublimation has apparently driven 2I’s activity from around 6 au inward (Fitzsimmons et al. 2019; Jewitt & Luu 2019; Ye et al. 2019), with detections of water products increasing toward perihelion. No evidence of water-ice rich grains was found on October 9 at 2.41 au by Yang et al. (2019). By October 11, the [OI] 6300 Å detection of McKay et al. (2019) implied, if it is produced solely by water dissociation, an H₂O production rate for 2I of $(6.3 \pm 1.5) \times 10^{26}$ mol s⁻¹. A marginal 18-cm OH production rate was measured from October 2–25 with the Nançay radio telescope at $(3.3 \pm 0.9) \times 10^{27}$ mol s⁻¹ (Crovisier et al. 2019). *Swift*/UVOT measurements show an increasing water production rate, going from $(7.0 \pm 1.5) \times 10^{26}$ to $(10.7 \pm 1.2) \times 10^{26}$ mol s⁻¹ from November 1 to December 1, 2019 (Xing et al. 2020).

We report the first observations from our ongoing detailed and comprehensive characterization of 2I’s composition and activity over time, using the Multi-Unit Spectroscopic Explorer (MUSE; Bacon et al. 2010) on the 8.2 m UT4 of the ESO/Very Large Telescope (VLT). Constraining the composition and nature of 2I’s volatiles and the production of its dust particles across the spatial distribution of the coma requires both spectroscopy and imaging (e.g. Opitom et al. 2019a). MUSE’s wide-field integral-field spectroscopy, with medium spectral resolution and sensitivity to faint and diffuse targets, is ideally suited to the task. Its 4800–9300 Å wavelength range contains several molecular and atomic species seen in emission in normal Solar System comets (e.g., Feldman et al. 2004). The species that directly leave the nucleus of the comet are termed “parent” species, and are non-detectable in the optical. They fragment to detectable “daughter” species, by many and varied paths, predominantly photodissociation. The common daughter species C₂ and CN have strong emission bands, with C₂ (0-0) at bandhead 5165 Å and the red CN system A²Π-X²Σ⁺, with bandheads at 9109 Å and 7822 Å. Several bands of the amino radical NH₂, the primary dissociation product of ammonia (NH₃), are also covered, such as the (0,3,0) band at 6300–6340 Å. Although forbidden oxygen is encompassed in the wavelength range (e.g. at 6363 Å), MUSE’s resolution is not sufficient to isolate it from its overlap with NH₂, and thus we do not place any constraint on H₂O production. Our future reports will provide simultaneous maps of the distribution in the coma of multiple gas species and dust, which we have sampled frequently throughout 2I’s perihelion passage.

In these November 2019 high-airmass and near-twilight observations, the slowly brightening comet is too faint for high-quality gas maps. Here we report detections, initial production rates and relative abundances of three gas species: C₂, NH₂, and CN. These species are commonly seen in the decades of measurements of Solar System comets (e.g. A’Hearn et al. 1995;

² JPL Horizons, 20 Nov 2019: <https://ssd.jpl.nasa.gov/sbdb.cgi?sstr=2I;old=0;orb=0;cov=0;log=0;cad=0#elem>

Fink 2009; Cochran et al. 2012; Hyland et al. 2019). Our results mark the first detection of NH_2 in 2I, and a definitive detection of C_2 in respect to previous searches (Fitzsimmons et al. 2019; Opitom et al. 2019b; Kareta et al. 2019; Lin et al. 2019). Observations by both of the TRAnsiting Planets and Planetesimals Small Telescope (TRAPPIST; Jehin et al. 2011), made simultaneously or near in time, provide supporting constraints on the violet CN and an upper limit on the C_2 production rates.

2. OBSERVATIONS AND DATA REDUCTION

2.1. MUSE

We observed 2I/Borisov with the MUSE integral-field spectrograph on VLT’s UT4 at Paranal, Chile, on 2019 November 14, 15, and 26, as detailed in Table 1. These measurements are the first part of our ongoing campaign to regularly monitor 2I through its perihelion passage³, providing a record of the comet’s composition and ejected dust, across a range of viewing angles and distances from the Sun. The coma of 2I has thus far been comparatively compact in our observations; the entire coma fits well within the 1×1 arcmin² MUSE WFM-mode field of view. MUSE’s FOV is sampled at $0.2''$ spatial resolution and operates with a spectral resolution of $R \simeq 3000$ at very high throughput (Bacon et al. 2010).

Our observing strategy at each epoch is based on previous successful observations of comets with MUSE (e.g. C/2015 ER61; Opitom et al. 2019a). We acquire four 600 s exposures, with VLT UT4 tracked at the comet’s on-sky rate of motion. We apply small dithering and a rotation of 90° between exposures. This allows later reduction of the signature of the detector gaps, streaked background stars, and better IFU-to-IFU uniformity in the combined data cubes. Sky exposures of 180 s were acquired, in an exposure sequence of OSOSO (O = on source, S = sky), for additional clean sky emission measurements. A spectrophotometric standard star was observed on the same night for later flux calibrations; LDS749B at airmass $X = 1.2$ on Nov 14, and CPD-69 177 at $X = 1.5$ on Nov 15, and at $X = 1.4$ on Nov 26. All observations were made in clear or photometric conditions, with $0.7\text{--}0.9''$ seeing. The November 14-15 observations had a high sky background from the nearly full Moon $80 - 90^\circ$ distant.

Dark subtraction, flat-fielding, telluric correction, and cube reconstruction of the MUSE observations were made with its data reduction pipeline⁴ (Weilbacher et al. 2016). Using the in-frame sky for simultaneous sky measurement was more accurate than sky subtraction from the sky exposures. The sky was thus estimated directly from zones covering about 20% of the science

exposures. For more details, we refer the reader to the MUSE pipeline manual. For the data from 14 and 15 November, we report only the first two exposures of each night. The final two exposures during twilight on 14 and 15 November had sufficiently strong sky background to be inadequately corrected by the pipeline, and we do not consider them further in this first analysis. The data cubes were flux calibrated and the telluric correction was made using the spectrophotometric standard star observations from that night. Because of the differential tracking and the lack of point-like sources in the field of view, the cubes obtained over the same night could not easily be re-combined using the MUSE pipeline. We thus reduced the cubes separately and combined the extracted spectra.

2.2. TRAPPIST

Near simultaneous observations were made with both 0.6-m TRAPPIST-North (TN) and TRAPPIST-South (TS) robotic telescopes located in the Atlas Mountains in Morocco and ESO/La Silla Observatory in Chile (Jehin et al. 2011). TS and TN are equipped with $2\text{k} \times 2\text{k}$ CCD cameras and HB comet narrow-band filters (Farnham et al. 2000) with a field of view of $\sim 20' \times 20'$. The images were binned 2×2 , resulting in a plate scale of $1.2''/\text{pixel}$. The narrow-band TS data for November 14–15 had low signal-to-noise, as the comet was at high airmass with full moon contamination, but good measurements were acquired with TN on November 10 and 17 with the CN filter. Data were acquired on November 25 with TN, and simultaneously with MUSE on November 26 with TS, in both cases under clear conditions and using both C_2 and CN filters (Table 1). The data reduction of the TRAPPIST narrow band images followed standard procedures using frequently updated master bias, flat, and dark frames. The dust contamination was removed, the sky contamination subtracted, and the flux calibration performed using regularly updated zero points based on observations of photometric standard stars, as described in Opitom et al. (2016) and Moulane et al. (2018).

3. ANALYSIS AND RESULTS

In Fig. 1, we present two MUSE spectra of the interstellar comet 2I/Borisov, each extracted over a $10''$ radius circular aperture. The spectra from 14 and 15 November are extracted from the two darkest-time cubes on each epoch, and subsequently co-added. Those from 26 November are from each of the four exposure cubes acquired that night, and then co-added. Cometary spectra are a superposition of the many emission bands from molecular gas, and the continuum of sunlight reflected from the dust in the coma. MUSE observations of 67P (where only dust was detected: Guilbert-Lepoutre et al. (2016)) were corrected for the difference in spectral slope between the two comets, and used to subtract the dust-reflected sunlight contribution, producing the spectra

³ ESO Director’s Discretionary program 2103.C-5070.

⁴ <http://www.eso.org/sci/software/pipelines/muse/muse-pipe-recipes.html>

Table 1. Observations of 2I/Borisov by MUSE/VLT and TRAPPIST

Time UT	Instrument/ Filter	Time MJD	Exposure (s)	Airmass	r_H (au)	Δ (au)	α (degrees)
2019-Nov-10 05:05	TN/CN	2458797.70833	1500	1.41	2.102	2.285	25.64
2019-Nov-14 08:03	MUSE	2458801.83542	600	2.07	2.078	2.230	26.26
2019-Nov-14 08:19	MUSE	2458801.84653	600	1.86	2.078	2.230	26.26
2019-Nov-15 08:02	MUSE	2458802.83472	600	2.04	2.072	2.217	26.40
2019-Nov-15 08:18	MUSE	2458802.84583	600	1.84	2.072	2.217	26.40
2019-Nov-17 05:00	TN/CN	2458804.70833	1500	1.45	2.061	2.191	26.65
2019-Nov-25 05:15	TN/C ₂	2458812.71885	1500	1.45	2.029	2.101	27.58
2019-Nov-25 05:40	TN/CN	2458812.72980	1500	1.38	2.029	2.101	27.58
2019-Nov-26 07:17	MUSE	2458813.80347	600	2.12	2.025	2.091	27.69
2019-Nov-26 07:34	MUSE	2458813.81528	600	1.89	2.025	2.091	27.69
2019-Nov-26 07:45	TS/C ₂	2458813.82291	1500	1.60	2.025	2.091	27.68
2019-Nov-26 07:46	MUSE	2458813.82361	600	1.76	2.025	2.091	27.69
2019-Nov-26 08:03	MUSE	2458813.83542	600	1.61	2.025	2.091	27.69
2019-Nov-26 08:10	TS/CN	2458813.84375	1500	1.44	2.024	2.090	27.69

NOTE— r_H and Δ are the heliocentric and geocentric distances, α the phase angle.

The MUSE November 14 observations were off-centred in the 1 arcmin² field of view, though contained all the apparent coma.

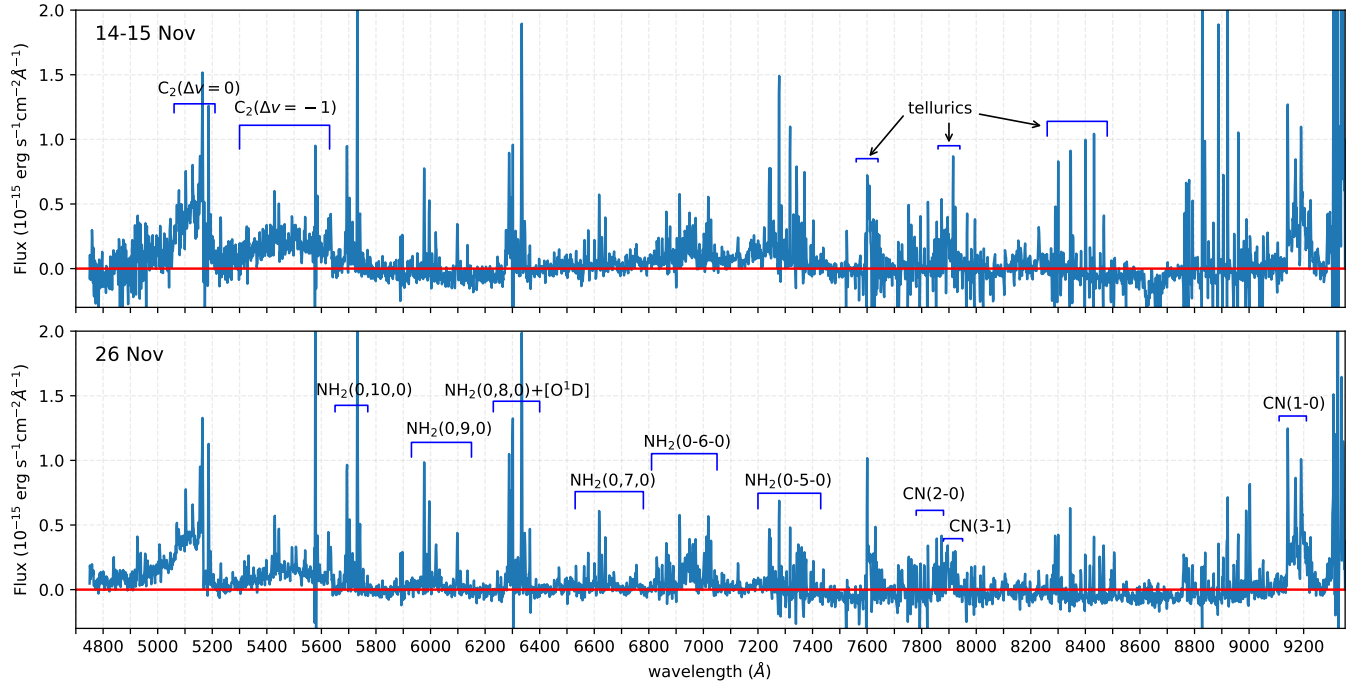


Figure 1. Spectra of the interstellar comet 2I/Borisov from two epochs: 14-15 November (combined due to bright moon conditions) and 26 November 2019, observed by the MUSE VLT integral-field spectrograph. On both co-added epochs, the full spectra are extracted over a 10'' radius aperture centred on the comet. Light reflected from 2I's dust has been subtracted using a spectral-slope-corrected continuum of 67P (Guilbert-Lepoutre et al. 2016). Several features are visible in both epochs, including the emission lines of the Swan system of C₂, many lines of NH₂ through the bands 10, 9, 8 and 7, and red CN. Gas species and some telluric emission lines, or terrestrial atmospheric lines, are indicated above the spectra. The majority of the non-labeled lines in the red region of the spectrum may be attributed to tellurics.

Table 2. Molecular fluorescence scattering efficiencies g and Haser model scale-lengths for parent and daughter species l_p , l_d at $r_H = 1$ au, used with $v = 0.5$ km/s

Band	g ergs s ⁻¹ mol ⁻¹	l_p km	l_d km
C ₂ ($\Delta v = 0$)	4.5×10^{-13}	2.2×10^4	6.6×10^4
NH ₂ (0-4-0)	9.2×10^{-15}	4.1×10^3	6.2×10^4
CN(1-0)	9.1×10^{-14}	1.3×10^4	2.1×10^5
CN(0-0)	2.6×10^{-13}	1.3×10^4	2.1×10^5

NOTE—All g -factors are assumed to scale as r_H^{-2} . All scalelengths are assumed to scale as r_H^2 , and are from A’Hearn et al. (1995) for C₂ and CN.

shown in Fig. 1. All of the most common Solar System comet gas species with emission bands in the MUSE wavelength coverage are present. The spectra exhibits the distinctive shape of C₂($\delta\nu = 0$) emission with its bandhead at 5165 Å, many notably bright red lines of NH₂, including those at 5733, 5976, 5995, 6021, 6334 Å (Cochran & Cochran 2002), the cometary [OI] forbidden lines overlapping with NH₂ at 6363Å, and the red CN (1–0) band around 9140 Å. Residuals from the sky subtraction in the red part of the spectrum are particularly apparent in the Nov 14–15 data, where the telluric O₂ band has been over-corrected.

The dust continuum in the MUSE spectra was fit using spectral regions free of significant gas-emission between 4750–9000 Å. We omit the measurement on November 14; the dust continuum shows large non-linear variations as a function of wavelength, possibly due to problems with sky subtraction due to the high airmass of observation, and we treat this measurement as unreliable. The spectral slope of the reddened continuum on November 26 was measured as 11.0 ± 0.2 %/10³ Å. This is similar to the spectral slope measured at red wavelengths by de León et al. (2019) on September 13, and so the dust colour of 2I appears consistent.

In Table 3, we present the flux and gas production rates in the coma of 2I for several nights spanning our observations in November 2019, and we consider each observed gas species in turn below. For the MUSE spectra, the gas species’ integrated flux was retrieved for each continuum-subtracted emission band (cf. Fig. 1). We then used the parameters given in Table 2 to convert the fluxes into column densities, as per the standard equation (e.g. Eq. 1, Beaver et al. 1990):

$$N = \frac{4\pi F}{\Omega g(r_H)} = 6.81 \times 10^{11} \frac{F r_H^2}{g(1\text{au})\theta^2} \quad (1)$$

where N is the column density, F the emission band’s flux, Ω the solid angle of the telescope’s aperture, r_H

the heliocentric distance, g the fluorescence scattering efficiency, and θ the diameter of the measurement aperture (here 10’). Using these column densities, we computed 2I’s gas production rates (Table 3) using a Haser model (Haser 1957). Note that while a few bands of CO⁺ are located in the MUSE spectral range, there is too much structure in the C₂ to reliably detect CO⁺ at present. Also, C₃ and N₂⁺ emission are both too far to the blue to see in MUSE data. At present, given the medium spectral resolution of MUSE, we do not identify any unusual/unidentified lines. For the TRAPPIST narrow-band images, the gas production rates were derived from median radial brightness profiles for the CN and C₂ images. The R broadband continuum filter images were used to remove the dust contamination. The fluxes were then converted to column densities and the profiles were fitted with a Haser model at 10,000 km from the nucleus (corresponding to 6.9 arcsec), using the same parameters as for the MUSE spectra (Table 2).

C₂ Swan ($\Delta v = 0$) Band-Sequence: In Fig. 2, we show the two detections of C₂ in the MUSE spectra. In order to interpret the results, we produced a model spectra assuming fluorescence equilibrium (Rousselot et al. 2000). Because the probability of intercombination transitions is low, this is a valid assumption, even if the timescale for C₂ radicals to equilibrate is long. This model provided a better fit to the measured spectrum than a simple mix of two Boltzmann populations. This is most likely due to the large 10 arcsec radius of the field of view used to extract these spectra, which would provide ample time for the majority of radicals to reach equilibrium. From this synthetic spectrum, we computed the whole flux received from the complete Swan $\Delta v = 0$ band, which are given in Table 3. The resulting production rate of $Q(\text{C}_2) = 1.1 \times 10^{24}$ mol s⁻¹ is above the upper limits from other facilities published previously; we discuss this further in § 4. There are other notable lines and features in the spectrum. For instance, at 5188 Å there is a feature due to the NH₂ radical. At 1 au, this line is usually much weaker than the C₂ feature.

NH₂ emission bands: Many bright emission lines of NH₂ are present in the MUSE spectra, including the AX (0,10,0)(0,0,0) and AX (0,8,0)(0,0,0). We fit the flux in the NH₂(0,10,0) band at 5680–5750 Å by fitting individual emission features within the band with gaussians and summing (Fig. 2), as the NH₂(0-8-0) 6324–6337 Å emission appeared too heavily contaminated by residuals from the sky subtraction of terrestrial [OI] and/or cometary [OI] emission. This is the first report of 2I’s NH₂ production rate.

CN(1-0) red emission band: We integrate the continuum-removed flux over 9100–9270 Å with IRAF as for the NH₂ band. This value was then corrected by a factor of 1.4 to account for the flux at longer wavelengths (Fink 2009), which were avoided because of the telluric noise, and potential second-order contamination.

Table 3. Flux from gas species detected in the coma of 2I/Borisov and their production rates in 2019 November

Date	Instrument	$C_2(\Delta v=0)$	$Q(C_2)$	$NH_2(0-10-0)$	$Q(NH_2)$	$CN(1-0)$	$CN(0-0)$	$Q(CN)$
			$\times 10^{24}$	$\times 10^{-15}$	$\times 10^{24}$	$\times 10^{-14}$	$\times 10^{-16}$	$\times 10^{24}$
		ergs $s^{-1}cm^{-2}$	mol s^{-1}	ergs $s^{-1}cm^{-2}$	mol s^{-1}	ergs $s^{-1}cm^{-2}$	ergs $s^{-1}cm^{-2}$	mol s^{-1}
Nov. 10	TN	6.6	1.9 ± 0.5
Nov. 14-15	MUSE	6.8×10^{-14}	1.1	1.4 ± 0.1	4.2	3.2 ± 0.3	...	1.8 ± 0.2
Nov. 17	TN	7.2	1.9 ± 0.5
Nov. 25	TN	$< 4.4 \times 10^{-16}$	< 2.0	6.2	1.6 ± 0.5
Nov. 26	MUSE	7.3×10^{-14}	1.1	1.5 ± 0.1	4.8	3.4 ± 0.3	...	1.8 ± 0.2
Nov. 26	TS	$< 4.3 \times 10^{-16}$	< 2.0	6.9	1.5 ± 0.5

NOTE—The measured C_2 and CN fluxes differ between TRAPPIST and MUSE because of the measurement technique. TN/TS fluxes are per arcsec² measured at a projected radial distance from the nucleus equivalent to 10,000 km at the comet. MUSE fluxes are total integrated fluxes through a circular aperture of 10'' radius.

The measurement uncertainties of $\sim 10\%$ are more substantial than any uncertainty in the fluorescence efficiency (g -factor) known for the CN red band; Paganini & Mumma (2016) find a g -factor of 41.9×10^{-3} photons/molecule/s at 1 au, which is very close to the Fink (1994) value of 41×10^{-3} photons/molecule/s.

CN(0-0) violet emission band: As the red emission band of CN is not as frequently used to measure CN production rates, we provide simultaneous or near-simultaneous constraints available from the TRAPPIST narrow-band imaging at 3890 Å. As shown in Table 3, the CN production rates are consistent between these independent measurements from the red and violet bands. This verifies that it is reasonable to compare the production rates of gas species in 2I measured from the red CN band. We note that Fink (e.g. 2009) also measured Solar System comets in the red band of CN, and had production rates consistent with those of other surveys such as A'Hearn et al. (1995) and Cochran et al. (2012), so there is precedent for using the red CN band to examine cometary composition.

4. DISCUSSION

2I's C_2 production has started in earnest (Fig. 3, upper left panel), given our high signal-to-noise prominent detection of C_2 (Fig. 2). In contrast, Fig. 3 (lower left panel) shows the CN production rate of 2I was stable within uncertainties between the end of September and the end of November. The higher CN production rate measurement on September 20 was under challenging conditions, since the sky background at this time was high; the SNR of Fitzsimmons et al. (2019) is much better than the marginal detection of Kareta et al. (2019). Records of the CN production rate at distances of several au are still sparse for dynamically new comets, though some have shown slow-increasing behaviour on inward approach (Whipple 1978).

The production of C_2 relative to that of CN has increased relatively quickly in the coma of 2I. Fig. 4 (left

panel) shows that the transition can be placed in the second week of November. It may be that a new source of C_2 , such as sublimation of organic-rich grains or another parent molecule, turned on once the comet came closer to the Sun. Cochran et al. (2012) found that C_2/CN production rate ratios are constant with heliocentric distance (see their Fig. 10). In contrast, Langland-Shula & Smith (2011) concluded from their much smaller population snapshot that the C_2/CN in the coma of comets tends to increase as the heliocentric distance diminishes, with a best-fit power-law index of 2.2 ± 0.6 . However, most of their measurements were within 2 au. Thus, the fast increase of 2I's production rate ratio may be notable: 2I was depleted in C_2 on October 13, in comparison to comets observed at the same heliocentric distance (Opitom et al. 2019b). A suitable comparison may be C/2009 P1 Garrard. This comet had a sudden increase in C_2 between 2.47 and 2.00 au inbound (McKay 2013), together with a huge change in CO/H₂O during its apparition (Feaga et al. 2014). Unlike 2I (Fitzsimmons et al. 2019), however, C/2009 P1 was unusual in that it was an exceptionally dusty comet. One interpretation of the change in composition of the coma is that the subsurface layers of 2I that are now contributing to its activity are less carbon-depleted than 2I's uppermost layer. In this scenario, inbound sublimation removed the C_2 -depleted layer. This would imply that our November observations of the coma displays a more representative volatile composition. It is also possible that a C_2 -rich region has been recently exposed, implying a heterogeneity of the nucleus.

There have been few rates of change of the C_2/CN production rates measured for individual Solar System comets outside 2 au. Langland-Shula & Smith (2011)'s four $Q(C_2)/Q(CN)$ measurements beyond 2 au are from only two different comets⁵. Intriguingly, Schulz et al.

⁵ 017P and C/1995 O1 (Hale-Bopp)

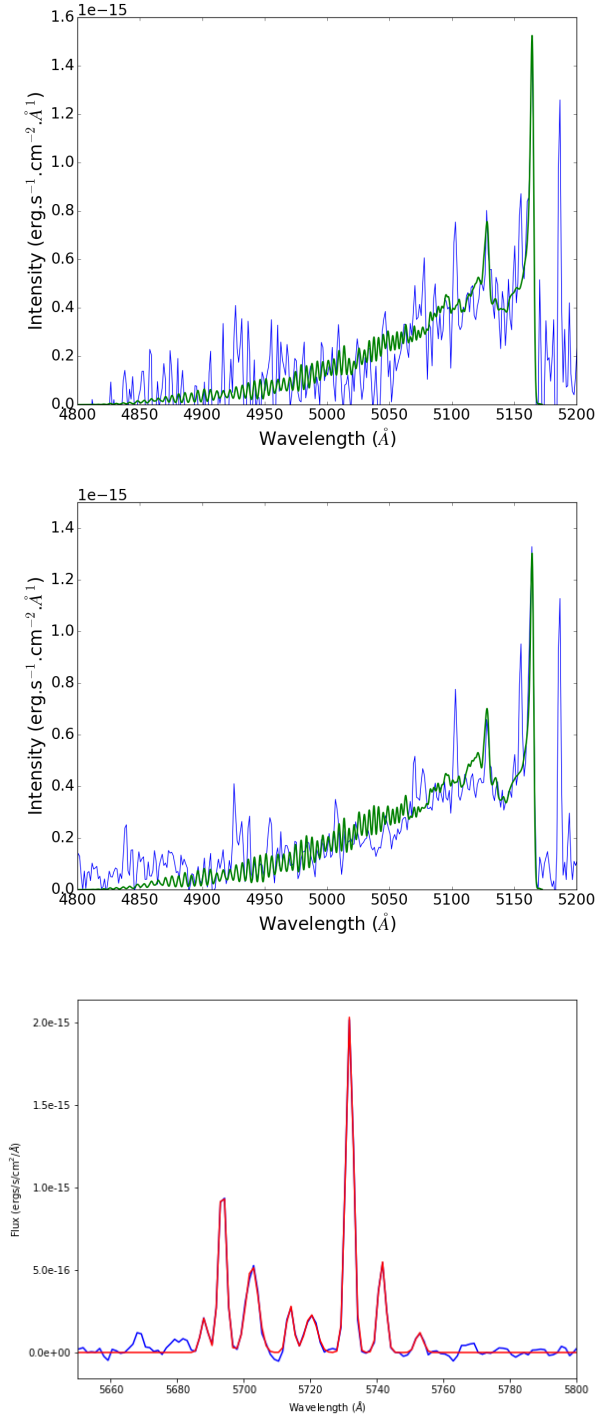


Figure 2. Detections in the MUSE spectra in Fig. 1: C_2 on 14-15 November (top), C_2 on 26 November (middle), and NH_2 on 26 November (lower). For C_2 , observational data are shown in blue, the green spectra are fits computed with a fluorescence equilibrium. The observed NH_2 (0,10,0) spectrum is shown in blue, with the gaussian fits to the component emission lines in red. Note that this is the NH_2 (0,3,0) band in modern notation as explained by (Cochran & Cochran 2002)

(1998) reported a strong evolution in $Q(C_2)/Q(CN)$ for comet 46P/Wirtanen, with the comet being C_2 depleted at 2 au, but having a normal abundance ratio at 1 au. However, Fink & Combi (2004) found this may be explained by varying Haser model parameters. Thus, it is hard to robustly say if 2I’s inbound abundance ratio change with r_H is unusually rapid. The lack of suitable Solar System comparison data reflects the challenge of measuring C_2 at heliocentric distances greater than a few au. Additionally, the parameters (cf. Table 2) and techniques used for computing production rates vary in the literature, and critically, these techniques are only resilient up to about 2 au. Acquiring and analysing comet gas production rates in a uniform manner, particularly in detail when comets are inbound between 3 and 2 au, will be a fruitful goal for 8m+ telescopes in the future — particularly given the expected increase in the rate of interstellar object discovery with the advent of the Rubin Observatory’s⁶ surveying (Moro-Martín et al. 2009; Cook et al. 2016; Trilling et al. 2017). We encourage the community to try to measure 2I’s gas production ratios as often as possible through and beyond 3 au, to see if the behaviour is symmetric.

Comparing the ratio of C_2 to CN production in 2I’s coma is useful to define its relationship to Solar System comet populations. Using a similar model and model parameters as we use in this work, A’Hearn et al. (1995) define a carbon-chain depleted comet as one with $\log[Q(C_2)/Q(CN)] < -0.18$. Based on their non-detections or weak detections of C_2 , previous studies concluded that 2I is depleted in carbon-chain species (Fig 4, right panel), similar to some Solar System comets (Kareta et al. 2019; Opitom et al. 2019b; Lin et al. 2019). According to A’Hearn et al. (1995)’s definition of carbon-chain depletion occurring at $\log[Q(C_2)/Q(CN)] \leq -0.21$, 2I is still carbon-chain depleted at the epochs we measure, but not nearly as much as was first thought (Fig. 4, right panel). In our Solar System, both Oort-family and Jupiter-family comets show carbon-chain depletion consistent with 2I’s depletion. Therefore, 2I is not directly analogous to any one dynamical family of comets. Unfortunately, it will not be possible to identify a likely region of origin in its star’s protoplanetary disk via comparison with the formation of Solar System comets.

The detection of NH_2 in the MUSE data is significant, because NH_2 is a dissociation product of NH_3 , which is a primary nitrogen reservoir in solar system comets. NH_2 is prominent in our spectra, yielding a ratio for 2I of $Q(NH_2)/Q(CN)=2.7$. Only a handful of optical studies of Solar System comet composition include measurements of the NH_2 production rates: Beaver et al. (1990) (6 comets at 0.65-1.8 au), Fink (2009) (50 comets), and Cochran et al. (2012) (26 comets

⁶ Formerly known as the Large Synoptic Survey Telescope.

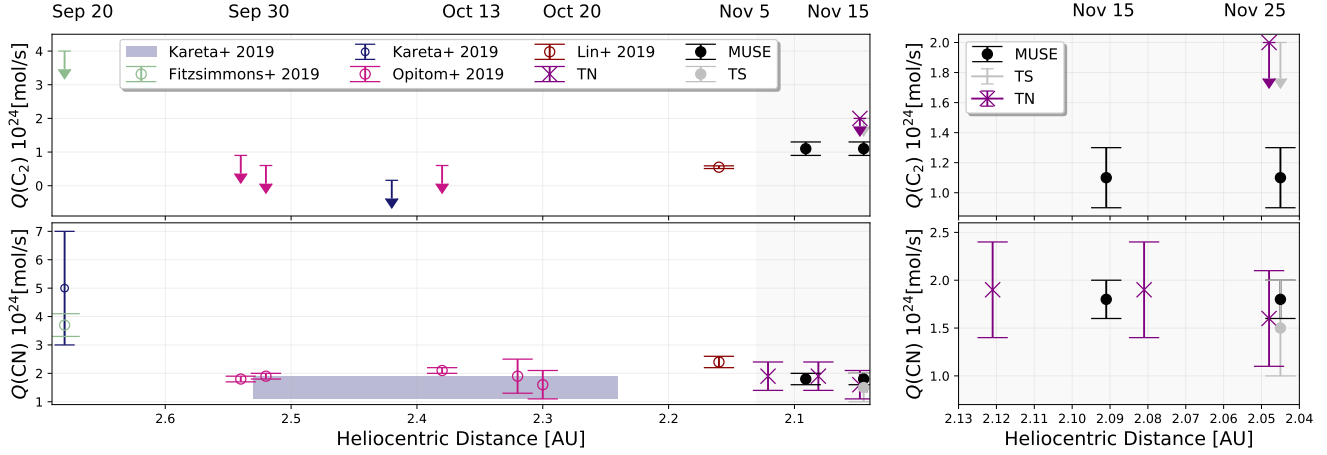


Figure 3. Pre-perihelion production rates of C_2 (upper) and CN (lower) in 2I/Borisov in 2019 September–November; its perihelion is 2019 December 8. The left two panels show the measurements and upper limits in [Fitzsimmons et al. \(2019\)](#), [Kareta et al. \(2019\)](#), [Opitom et al. \(2019b\)](#), [Lin et al. \(2019\)](#), and this work. The dark blue shaded region in the bottom left panel shows the continuous upper and lower limit measurements provided in [Kareta et al. \(2019\)](#) for the production of CN during this period. The right two panels show our observations with MUSE, Trappist North and Trappist South.

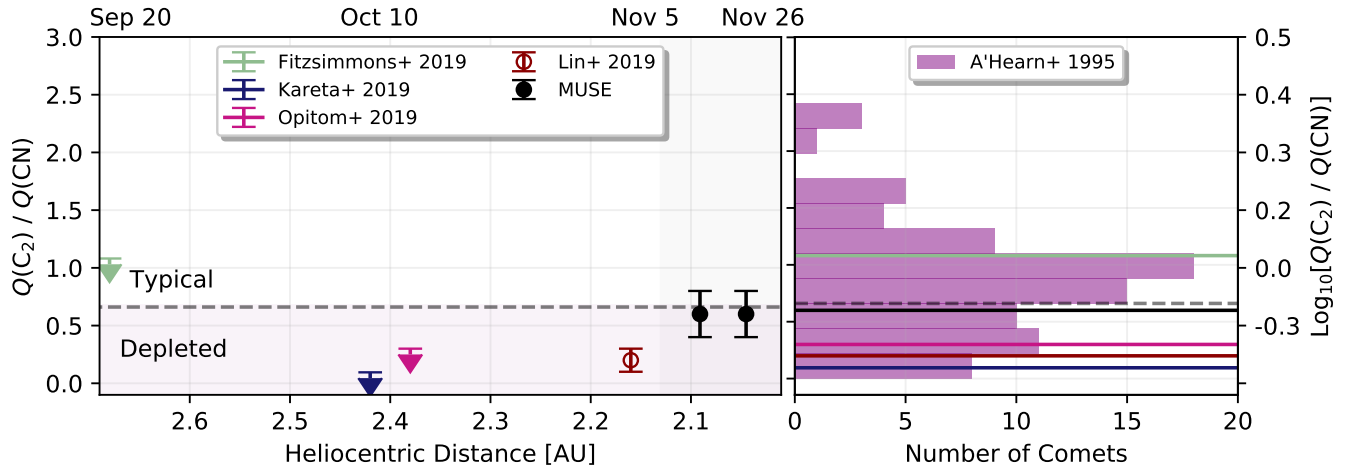


Figure 4. Pre-perihelion ratio of the production rates of C_2 and CN in 2I/Borisov in 2019 September–November; its perihelion is 2019 December 8. We show the measurements and upper limits in [Fitzsimmons et al. \(2019\)](#), [Kareta et al. \(2019\)](#), [Opitom et al. \(2019b\)](#), [Lin et al. \(2019\)](#), and this work as a function of 2I/Borisov’s heliocentric distance in the left panel. The two black points show the ratios measured with the observations with MUSE on November 14–15 and November 25, when both CN and C_2 are detected at high SNR. In the right panel, the purple histograms show the ratio of production rates for comets in the sample presented in [A’Hearn et al. \(1995\)](#). The solid lines indicate the measured values and upper limits for Borisov, with the same color scheme as in the left panel. The black dashed line indicates the commonly accepted definition for carbon depleted comets, $\log[Q(C_2)/Q(CN)] \leq -0.18$

with NH_2 detections that were observed on at least three nights, although with some overlap with the sample of [Fink \(2009\)](#)). Fig. 5 shows that the NH_2 to CN production ratio that we measure for 2I at 2.02 au of $\log[Q(NH_2)/Q(CN)] = 0.43$ is higher than almost all of the comets in the [Cochran et al. \(2012\)](#) and [Fink \(2009\)](#) data sets, with the exception only of comets 102P and C/1988 Y1. [Cochran et al. \(2012\)](#) report a mean log production rate ratio with respect to CN of 0.09 ± 0.25

for the 26 comets with NH_2 in their restricted dataset; thus 2I is within 1.4σ . While there is no notable trend of the NH_2/CN production ratio currently known between 1 and 2 au, we note that for [Fink \(2009\)](#)’s data set, only two comets have NH_2 production rates measured beyond 2 au (one is relatively rich and one is poor in NH_2), and for [Cochran et al. \(2012\)](#), there are no values of the NH_2/CN ratio beyond 2 au. The current statistics are insufficient to refine the comparison

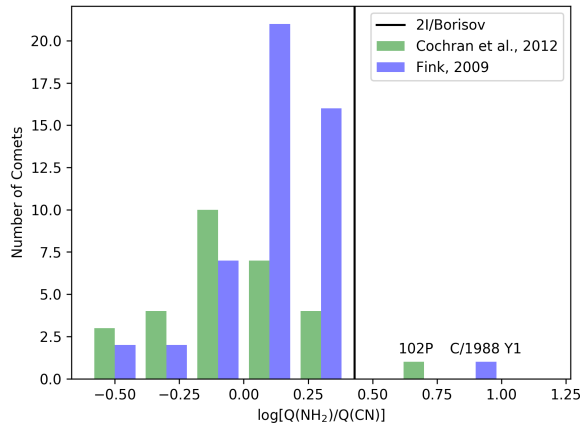


Figure 5. Ratio of the production rates of NH₂ and CN for the comets with NH₂ detections reported by Fink (2009); Cochran et al. (2012). The bin size is larger than the measurement uncertainties, where they exist. Only two comets (labelled) have a higher ratio than 2I/BorISOV.

of 2I’s NH₂/CN production ratio more finely relative to a dynamical origin as Oort or Jupiter family comets. Comparison to literature values requires care, because the MUSE spectrum has been extracted close to the nucleus (within 10 arcsec), and NH₂ has a much shorter lifetime than C₂ and CN, while usually long-slit spectra are extracted on a larger area (~ 1 arcmin). Nuclei size should also be considered, as at 0.2–0.5 km, 2I is smaller than most well-studied comets to date. Nevertheless, 2I is relatively rich in NH₂ compared to the available bulk of Solar System comets.

The kinship of 2I’s physical properties implies an origin system that had a planetesimal-forming disk with many chemical similarities to our own. Many difficulties remain in linking the span of processes from disk to comets. Our own system’s cometary populations appear largely homogeneous, with subtle compositional variation (c.f. Cochran et al. 2012). Identifying nuanced differences that could produce variation like the NH₂ enrichment we observe in 2I will be an ongoing task.

5. CONCLUSION

We report the detection of C₂, NH₂ and the red bandhead of CN in 2I/BorISOV, the first interstellar object detected passing through the Solar System with prominent cometary activity. We present measurements from the MUSE ESO/VLT integral-field spectrograph taken on three nights in November 2019. These measurements are the first of a larger observing program that is monitor-

ing 2I through its perihelion passage. We combine these data with complementary observations from the two TRAPPIST telescopes in order to constrain the production rates of various species in the coma of 2I. The CN production rates measured near-simultaneously from the violet and red bandheads of CN emission are consistent within uncertainties. As of 2019 November 26, the dust continuum of 2I has a spectral slope in the optical of $11.0 \pm 0.2 \text{ \%}/10^3 \text{ \AA}$, consistent with earlier reports. The C₂ production of 2I has activated, and a changing C₂/CN production ratio is apparent. At this epoch, 2I is only slightly depleted in C₂ at $Q(\text{C}_2)/Q(\text{CN})=0.61$, but it is rich in NH₂ relative to almost all Solar System comets, at $Q(\text{NH}_2)/Q(\text{CN})=2.7$. The known sample of Solar System comets observed at 2 au and beyond is sufficiently small to hamper comparison, particularly in the production rates for NH₂ relative to that of CN in carbon-chain depleted comets. 2I’s chemical composition retains many similarities to the broader population of Solar System comets.

ACKNOWLEDGMENTS

M.T.B. and A.F. appreciate support from UK STFC grant ST/P0003094/1. M.M.K. was supported by NASA Solar System Observations grant 80NSSC18K0856. D.S. is supported by the National Aeronautics and Space Administration through the NASA Astrobiology Institute under Cooperative Agreement Notice NNH13ZDA017C issued through the Science Mission Directorate. M.M. was supported by the National Aeronautics and Space Administration under Grant No. 80NSSC18K0849 issued through the Planetary Astronomy Program. A.G.L. was supported by the European Research Council (ERC) under grant agreement No. 802699.

Based on observations collected at the European Southern Observatory under ESO programme 2103.C-5070. We thank the ESO staff, particularly Henri Boffin, Bin Yang, Diego Parraguez, Edmund Christian Herenz, Fuyan Bian, and Israel Blanchard, for their help in the acquisition of these observations.

TRAPPIST is a project funded by the Belgian Fonds (National) de la Recherche Scientifique (F.R.S.-FNRS) under grant FRFC 2.5.594.09.F and the ARC grant for Concerted Research Actions, financed by the Wallonia-Brussels Federation. TRAPPIST-North is a project funded by the University of Liege, in collaboration with Cadi Ayyad University of Marrakech (Morocco). E.J is a F.R.S.-FNRS Senior Research Associate.

REFERENCES

- A’Hearn, M. F., Millis, R. C., Schleicher, D. O., Osip, D. J., & Birch, P. V. 1995, *Icarus*, 118, 223, doi: [10.1006/icar.1995.1190](https://doi.org/10.1006/icar.1995.1190)
- Bacon, R., Accardo, M., Adjali, L., et al. 2010, Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, Vol. 7735, The MUSE second-generation VLT instrument, 773508, doi: [10.1117/12.856027](https://doi.org/10.1117/12.856027)

- Beaver, J. E., Wagner, R. M., Schleicher, D. G., & Lutz, B. L. 1990, *ApJ*, 360, 696, doi: [10.1086/169155](https://doi.org/10.1086/169155)
- Bolin, B. T., Lisse, C. M., Kasliwal, M. M., et al. 2019, arXiv e-prints, arXiv:1910.14004. <https://arxiv.org/abs/1910.14004>
- Cochran, A. L., Barker, E. S., & Gray, C. L. 2012, *Icarus*, 218, 144, doi: [10.1016/j.icarus.2011.12.010](https://doi.org/10.1016/j.icarus.2011.12.010)
- Cochran, A. L., & Cochran, W. D. 2002, *Icarus*, 157, 297, doi: [10.1006/icar.2002.6850](https://doi.org/10.1006/icar.2002.6850)
- Cook, N. V., Ragozzine, D., Granvik, M., & Stephens, D. C. 2016, *ApJ*, 825, 51, doi: [10.3847/0004-637X/825/1/51](https://doi.org/10.3847/0004-637X/825/1/51)
- Crovisier, J., Colom, P., B. N., & Bockelée-Morvan, D. 2019, CBET 4691, International Astronomical Union. <http://www.cbat.eps.harvard.edu/iau/cbet/004600/CBET004691.txt>
- de León, J., Licandro, J., Serra-Ricart, M., et al. 2019, *Research Notes of the American Astronomical Society*, 3, 131, doi: [10.3847/2515-5172/ab449c](https://doi.org/10.3847/2515-5172/ab449c)
- Engelhardt, T., Jedicke, R., Vereš, P., et al. 2017, *AJ*, 153, 133, doi: [10.3847/1538-3881/aa5c8a](https://doi.org/10.3847/1538-3881/aa5c8a)
- Farnham, T. L., Schleicher, D. G., & A’Hearn, M. F. 2000, *Icarus*, 147, 180, doi: [10.1006/icar.2000.6420](https://doi.org/10.1006/icar.2000.6420)
- Feaga, L. M., A’Hearn, M. F., Farnham, T. L., et al. 2014, *AJ*, 147, 24, doi: [10.1088/0004-6256/147/1/24](https://doi.org/10.1088/0004-6256/147/1/24)
- Feldman, P. D., Cochran, A. L., & Combi, M. R. 2004, *Spectroscopic investigations of fragment species in the coma*, ed. M. C. Festou, H. U. Keller, & H. A. Weaver, 425
- Fink, U. 1992, *Science*, 257, 1926, doi: [10.1126/science.257.5078.1926](https://doi.org/10.1126/science.257.5078.1926)
- . 1994, *ApJ*, 423, 461, doi: [10.1086/173824](https://doi.org/10.1086/173824)
- . 2009, *Icarus*, 201, 311, doi: [10.1016/j.icarus.2008.12.044](https://doi.org/10.1016/j.icarus.2008.12.044)
- Fink, U., & Combi, M. R. 2004, *Planet. Space Sci.*, 52, 573, doi: [10.1016/j.pss.2003.12.002](https://doi.org/10.1016/j.pss.2003.12.002)
- Fitzsimmons, A., Hainaut, O., Meech, K., et al. 2019, In press, *ApJL*, arXiv:1909.12144. <https://arxiv.org/abs/1909.12144>
- Guilbert-Lepoutre, A., Besse, S., Snodgrass, C., & Yang, B. 2016, in *AAS/Division for Planetary Sciences Meeting Abstracts #48*, AAS/Division for Planetary Sciences Meeting Abstracts, 116.12
- Guzik, P., Drahus, M., Rusek, K., et al. 2019, *Nature Astronomy*, 467, doi: [10.1038/s41550-019-0931-8](https://doi.org/10.1038/s41550-019-0931-8)
- Haser, L. 1957, *Bulletin de la Societe Royale des Sciences de Liege*, 43, 740
- Hyland, M. G., Fitzsimmons, A., & Snodgrass, C. 2019, *MNRAS*, 484, 1347, doi: [10.1093/mnras/stz075](https://doi.org/10.1093/mnras/stz075)
- Jehin, E., Gillon, M., Queloz, D., et al. 2011, *The Messenger*, 145, 2
- Jewitt, D., Hui, M.-T., Kim, Y., et al. 2019, arXiv e-prints, arXiv:1912.05422. <https://arxiv.org/abs/1912.05422>
- Jewitt, D., & Luu, J. 2019, arXiv e-prints, arXiv:1910.02547. <https://arxiv.org/abs/1910.02547>
- Kareta, T., Andrews, J., Noonan, J. W., et al. 2019, arXiv e-prints, arXiv:1910.03222. <https://arxiv.org/abs/1910.03222>
- Langland-Shula, L. E., & Smith, G. H. 2011, *Icarus*, 213, 280, doi: [10.1016/j.icarus.2011.02.007](https://doi.org/10.1016/j.icarus.2011.02.007)
- Lee, C.-H., Lin, H.-W., Chen, Y.-T., & Yen, S.-F. 2019, *Research Notes of the American Astronomical Society*, 3, 184, doi: [10.3847/2515-5172/ab5f69](https://doi.org/10.3847/2515-5172/ab5f69)
- Lin, H. W., Lee, C.-H., Gerdes, D. W., et al. 2019, arXiv e-prints, arXiv:1912.06161. <https://arxiv.org/abs/1912.06161>
- McKay, A. J. 2013, PhD thesis, New Mexico State University
- McKay, A. J., Cochran, A. L., Dello Russo, N., & DiSanti, M. 2019, arXiv e-prints, arXiv:1910.12785. <https://arxiv.org/abs/1910.12785>
- Meech, K. J., Weryk, R., Micheli, M., et al. 2017, *Nature*, 552, 378, doi: [10.1038/nature25020](https://doi.org/10.1038/nature25020)
- Micheli, M., Farnocchia, D., Meech, K. J., et al. 2018, *Nature*, 559, 223, doi: [10.1038/s41586-018-0254-4](https://doi.org/10.1038/s41586-018-0254-4)
- Moro-Martín, A., Turner, E. L., & Loeb, A. 2009, *ApJ*, 704, 733, doi: [10.1088/0004-637X/704/1/733](https://doi.org/10.1088/0004-637X/704/1/733)
- Moulane, Y., Jehin, E., Opitom, C., et al. 2018, *Astronomy & Astrophysics*, 619, A156, doi: [10.1051/0004-6361/201833582](https://doi.org/10.1051/0004-6361/201833582)
- Opitom, C., Yang, B., Selman, F., & Reyes, C. 2019a, *A&A*, 628, A128, doi: [10.1051/0004-6361/201833960](https://doi.org/10.1051/0004-6361/201833960)
- Opitom, C., Guilbert-Lepoutre, A., Jehin, E., et al. 2016, *A&A*, 589, A8, doi: [10.1051/0004-6361/201527628](https://doi.org/10.1051/0004-6361/201527628)
- Opitom, C., Fitzsimmons, A., Jehin, E., et al. 2019b, *A&A*, 631, L8, doi: [10.1051/0004-6361/201936959](https://doi.org/10.1051/0004-6361/201936959)
- ‘Oumuamua ISSI Team, Bannister, M. T., Bhandare, A., et al. 2019, *Nature Astronomy*, 3, 594, doi: [10.1038/s41550-019-0816-x](https://doi.org/10.1038/s41550-019-0816-x)
- Paganini, L., & Mumma, M. J. 2016, *ApJS*, 226, 3, doi: [10.3847/0067-0049/226/1/3](https://doi.org/10.3847/0067-0049/226/1/3)
- Rousselot, P., Hill, S. M., Burger, M. H., et al. 2000, *Icarus*, 146, 263, doi: [10.1006/icar.2000.6383](https://doi.org/10.1006/icar.2000.6383)
- Schleicher, D. G. 2008, *AJ*, 136, 2204, doi: [10.1088/0004-6256/136/5/2204](https://doi.org/10.1088/0004-6256/136/5/2204)
- Schulz, R., Arpigny, C., Manfroid, J., et al. 1998, *A&A*, 335, L46
- Trilling, D. E., Robinson, T., Roegge, A., et al. 2017, *ApJL*, 850, L38, doi: [10.3847/2041-8213/aa9989](https://doi.org/10.3847/2041-8213/aa9989)

Weilbacher, P. M., Streicher, O., & Palsa, R. 2016,

MUSE-DRP: MUSE Data Reduction Pipeline.

<http://ascl.net/1610.004>

Whipple, F. L. 1978, *Moon and Planets*, 18, 343,

doi: [10.1007/BF00896489](https://doi.org/10.1007/BF00896489)

Xing, Z., Bodewits, D., Noonan, J., & Bannister, M. T.

2020, arXiv e-prints, arXiv:2001.04865.

<https://arxiv.org/abs/2001.04865>

Yang, B., Kelley, M. S. P., Meech, K. J., et al. 2019, arXiv e-prints, arXiv:1912.05318.

<https://arxiv.org/abs/1912.05318>

Ye, Q., Kelley, M. S. P., Bolin, B. T., et al. 2019, arXiv e-prints, arXiv:1911.05902.

<https://arxiv.org/abs/1911.05902>