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Astrochemical relevance of VUV ionization of large PAH cations*

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ABSTRACT

Context. As a part of interstellar dust, polycyclic aromatic hydrocarbons (PAHs) are processed by the interaction with vacuum ultraviolet (VUV) photons that are emitted by hot young stars. This interaction leads to the emission of the well-known aromatic infrared bands but also of electrons, which can significantly contribute to the heating of the interstellar gas.

Aims. Our aim is to investigate the impact of molecular size on the photoionization properties of cationic PAHs.

Methods. Trapped PAH cations of sizes between 30 and 48 carbon atoms were submitted to VUV photons in the range of 9 to 20 eV from the DESIRS beamline at the synchrotron SOLEIL. All resulting photoproducts including dications and fragment cations were mass-analyzed and recorded as a function of photon energy.

Results. Photoionization is found to be predominant over dissociation at all energies, which differs from an earlier study on smaller PAHs. The photoionization branching ratio reaches 0.98 at 20 eV for the largest studied PAH. The photoionization threshold is observed to be between 9.1 and 10.2 eV, in agreement with the evolution of the ionization potential with size. Ionization cross sections were indirectly obtained and photoionization yields extracted from their ratio with theoretical photoabsorption cross sections, which were calculated using time-dependent density functional theory. An analytical function was derived to calculate this yield for a given molecular size.

Conclusions. Large PAH cations could be efficiently ionized in H I regions and provide a contribution to the heating of the gas by photoelectric effect. Also, at the border of or in H II regions, PAHs could be exposed to photons of energy higher than 13.6 eV. Our work provides recipes to be used in astronomical models to quantify these points.

Key words. astrochemistry – methods: laboratory: molecular – molecular processes – ISM: molecules – ISM: dust, extinction – ultraviolet: ISM

1. Introduction

Polycyclic aromatic hydrocarbons (PAHs) play a major role in the physics and chemistry of photodissociation regions (PDRs). They strongly absorb vacuum ultraviolet (VUV) photons that are emitted by hot young stars and relax by emission in the aromatic infrared bands (AIBs). The interaction with VUV photons can lead to other relaxation processes including ionization and dissociation. All these processes together with reactive processes involving in particular electrons and hydrogen (H, H₂) govern the evolution of the PAH population in the diffuse interstellar medium (Le Page et al. 2003), in circumstellar disks (Visser et al. 2007), and in reflection nebulae (Montillaud et al. 2013). The results of these chemical models suggest that large PAHs with

a typical carbon number, N_C , of 50 or more dominate the AIB emission which led to the grandPAH hypothesis that large and possibly compact PAHs dominate the emission in bright PDRs (Andrews et al. 2015). In some regions associated with these PDRs, large PAHs are expected to be ionized reaching even the dicationic stage (Tielens 2005; Andrews et al. 2016).

In a previous study we have investigated the branching ratio between ionization and fragmentation upon VUV irradiation for medium-sized PAH cations (Zhen et al. 2016) with an N_C between 16 and 24. For all of these cations, fragmentation was observed to be the dominant channel at least up to a photon energy of 13.6 eV which is relevant for H I regions. In the case of larger PAH cations, ionization is expected to be by far the dominant channel as suggested by the study of the hexa-peri-hexabenzocoronene (HBC) cation, C₄₂H₁₈⁺, by Zhen et al. (2015). Here, our objective is to quantify the growing importance of ionization as the molecular size increases. Follow-

* The dataset associated with this work can be found under 10.5281/zenodo.3899775

† Corresponding author.

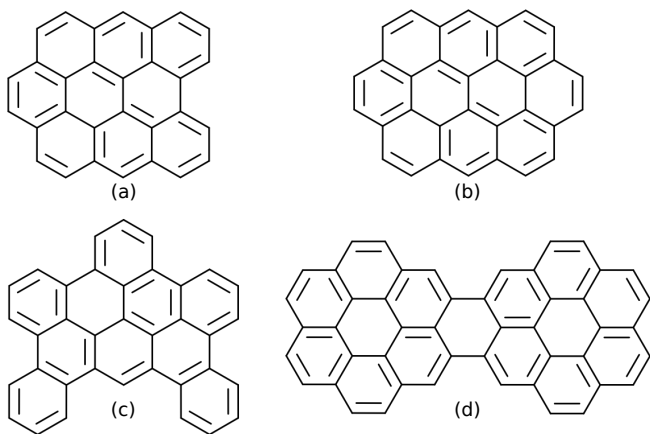


Fig. 1. Molecular structures of the studied PAHs, namely (a) benzo[ghi]perylene, $C_{30}H_{14}^+$, (b) ovalene, $C_{32}H_{14}^+$, (c) DBPP, $C_{36}H_{18}^+$, and (d) dicoronylene, $C_{48}H_{20}^+$.

ing Zhen et al. (2016), we have studied the photoprocessing of PAH cations with an N_C between 30 and 48 atoms over the 9.5 – 20.0 eV VUV range. Photon energies above the Lyman limit are relevant to PAHs observed at the border of ionization fronts in PDRs (Vicente et al. 2013), as well as in H II regions (Compiègne et al. 2007).

Although we report here also the branching ratio between ionization and dissociation, our analysis is focused on ionization. More specifically, we derive the photoionization yield, which is important to model the charge balance of PAHs and its impact on the AIB spectrum (Bakes et al. 2001), but also to evaluate the contribution of these species to the photoelectric heating rate (Bakes & Tielens 1994; Weingartner & Draine 2001b). The experimental method is described in Sect. 2 and the results are presented in Sect. 3. In Sect. 4, we discuss the astrophysical implications and propose recipes to be used in astrophysical models. We conclude in Sect. 5.

2. Experimental method and data analysis

We have used the Thermo Scientific™ LTQ XL™ linear ion trap (LTQ ion trap) as described in Milosavljević et al. (2012), which is available at the VUV beamline DESIRS at the synchrotron SOLEIL (Nahon et al. 2012).

The production of PAH cations in the LTQ ion trap was performed using an atmospheric pressure photoionization (APPI) source which required the species of interest to be in solution before their injection with a syringe. This part was a major limitation on the size range of PAHs we could study due to the non-solubility of large PAHs. Four large PAH cations with N_C ranging from 30 to 48 could be investigated in this study, namely (a) benzo[ghi]perylene, $C_{30}H_{14}^+$, (b) ovalene, $C_{32}H_{14}^+$, (c) dibenzophenanthrene (DBPP), $C_{36}H_{18}^+$, and (d) dicoronylene, $C_{48}H_{20}^+$. Sample (b) originated from Janssen Chimica (Belgium), samples (a) and (c) from the PAH Research Institute in Greifenberg (Dr. Werner Schmidt). The synthesis of compound (d) is briefly reported in Appendix A. The molecular structures of the studied species are depicted in Fig. 1. Emptying the syringe was performed at a flow rate which was kept constant for each experiment. The used flow rate was $4 \mu\text{l min}^{-1}$ for compounds (a) and (b), $6 \mu\text{l min}^{-1}$ for compound (c), and $10 \mu\text{l min}^{-1}$ for compound (d). The presence of UV irradiation from a Kr discharge lamp ensured a soft creation of PAH cations without fragmentation (Giuliani et al. 2012). The formed cations were then guided

through ion optics into the LTQ ion trap in which a constant He pressure of $p \approx 10^{-3}$ mbar was held. The ions were cooled by the collisions with He atoms and the PAH cations of interest, the so-called parent ions, were isolated through specific mass selection and ejection of other species from the ion trap including the ^{13}C isotopomers.

The parent ions were then submitted to the VUV synchrotron radiation which was tuned from 9.5 to 20.0 eV in steps of 0.1, 0.2, 0.3, or 0.5 eV depending on the photon energy range, with the exception of $C_{30}H_{14}^+$, for which we were able to scan only at low energies up to 15.5 eV. Higher harmonics of the VUV undulator synchrotron radiation with photon energies lower than 16.0 eV were filtered out by a gas filter filled with Ar gas to a pressure of 0.23 mbar. Above 16.0 eV no such gas filtering is necessary. The photon flux was measured with an IRD AXUV100 calibrated Si photodiode for a monochromator exit slit width of $200 \mu\text{m}$ and was between 0.8 and $2.8 \cdot 10^{12}$ photons s^{-1} over the studied photon energy range. A typical photon flux can be derived using a previous calibration of the beam size as a function of photon energy (Douix et al. 2017), yielding values of $1.5 - 5.2 \cdot 10^{14}$ photons $\text{cm}^{-2}\text{s}^{-1}$. In order to limit possible two photon consecutive absorption processes, we tuned the photon flux by changing (i) the irradiation time from 0.8 to 0.2 s for the lower and higher photon energy ranges, respectively, and (ii) the monochromator exit slit width from $200 \mu\text{m}$ at low energies to $70 \mu\text{m}$ at high energies, except for dicoronylene for which values of 400 and $100 \mu\text{m}$ at low and high energies, respectively, were used to improve the signal-to-noise ratio. The photon dose was assumed to be linearly proportional to both the irradiation time and the monochromator exit slit width. The probability of two photon absorption processes could be estimated on the formation of triply charged parent ions, yielding only very small relative intensities below 2 % of the total number of photoproducts.

Depending on the acquisition time, a few hundred mass spectra were recorded at each photon energy and averaged to yield one mass spectrum per photon energy step. Following the same procedure, we also recorded blank mass spectra at each photon energy by selecting a mass close to but different enough from each parent ion. This allowed us to perform background subtraction which eliminates contamination peaks from the mass spectra. The averaging procedure provides us with a statistical standard error (see Appendix D). As an example, the background subtracted mass spectra for ovalene, which has a mass-to-charge ratio of $m/z = 398$, are depicted in Fig. 2 at two different photon energies of 9.5 eV and 15.5 eV. The parent ion, $C_{32}H_{14}^+$, is well isolated, the ^{13}C isotopic parent ion has a residual contribution of less than 1 % remaining in the ion trap. By increasing the photon energy, three different secondary ions can be observed and unambiguously separated, namely the H and 2H/H₂ loss, and the main doubly ionized parent ion channels. For the presented example of the ovalene cation, $C_{32}H_{14}^+$, these species are $C_{32}H_{13}^+$, $C_{32}H_{12}^+$, and $C_{32}H_{14}^{2+}$, respectively (see Fig. 2). When extracting the peak intensities as will be done in the following, one has to consider the detector gain efficiency that varies with the charge and the mass of the ions of interest. Recommended scaling factors were therefore applied (see Appendix B).

3. Results and discussion

3.1. Action spectra and branching ratio

The action spectra are determined following the procedure described in Appendix B yielding relative intensities of the pho-

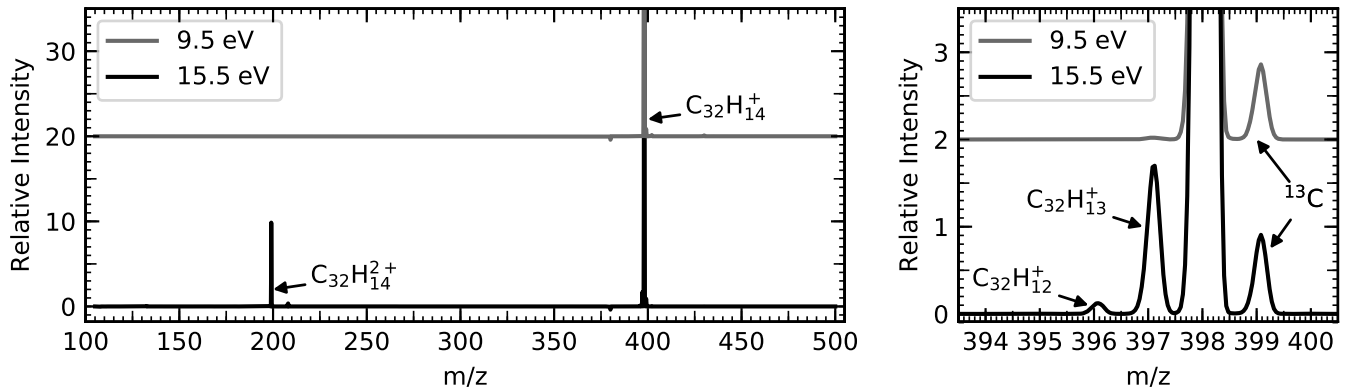


Fig. 2. Mass spectra of the ovalene parent cation, $C_{32}H_{14}^+$, $m/z = 398$, at two different photon energies. At 9.5 eV, none of the photoionization or photodissociation channels are opened whereas at 15.5 eV, the doubly ionized parent ion, $C_{32}H_{14}^{2+}$, at $m/z = 199$ is observed as well as the photofragments due to the loss of one and two hydrogen atoms, $C_{32}H_{13}^+$ and $C_{32}H_{12}^+$ at $m/z = 397$ and 396, respectively.

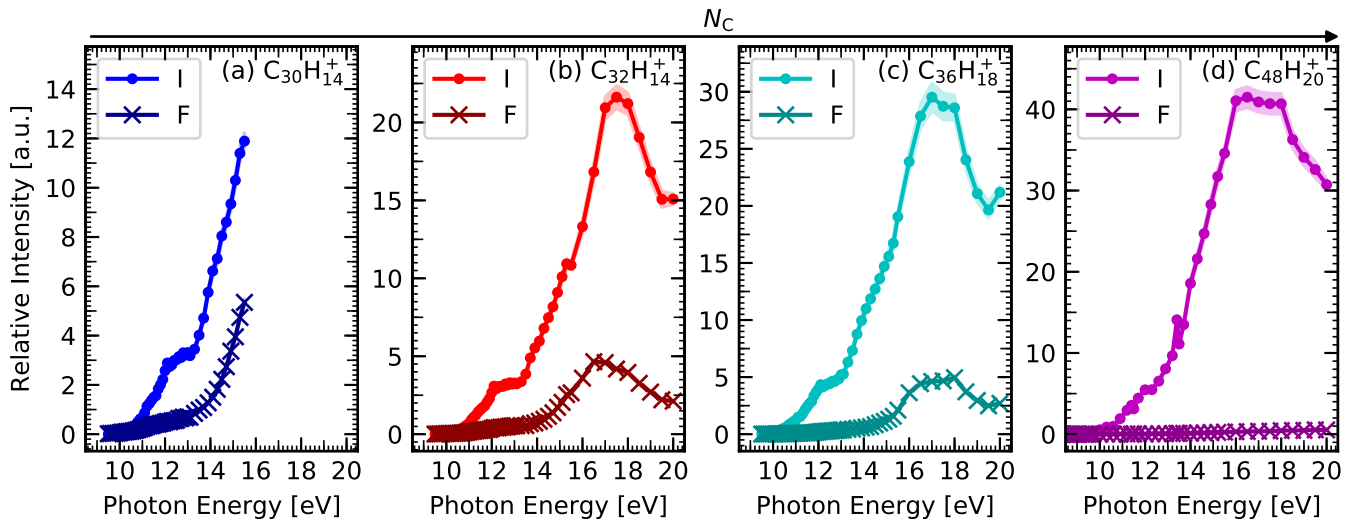


Fig. 3. Action spectra of the photoproducts, dications (I) and fragments (F), as a function of photon energy for all studied PAH cations, (a) benzobisanthrene, (b) ovalene, (c) DBPP, and (d) dicoronylene after absorption of a VUV photon. Relative intensities as explained in Appendix B.

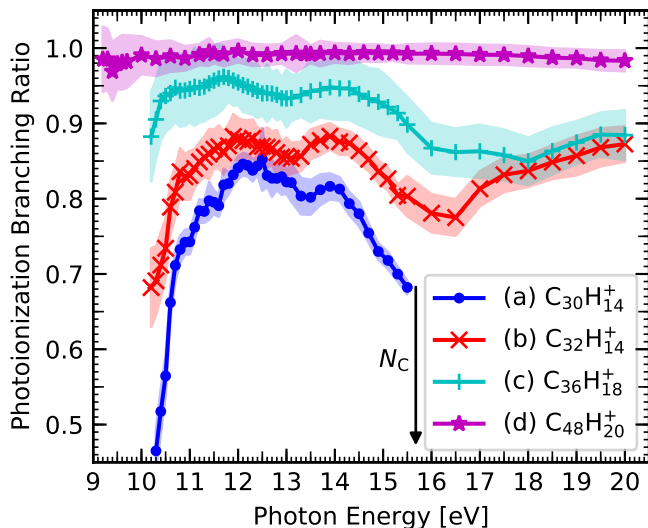


Fig. 4. Photoionization branching ratio relative to photodissociation as a function of photon energy, for (a) benzobisanthrene, $C_{30}H_{14}^+$, (b) ovalene, $C_{32}H_{14}^+$, (c) DBPP, $C_{36}H_{18}^+$, and (d) dicoronylene, $C_{48}H_{20}^+$, after absorption of a single VUV photon.

toproducts as used in previous work (Zhen et al. 2016). The resulting spectra for the photoionization (dication, denoted I) and photodissociation (fragments, denoted F) channels of the four studied PAH cations are shown in Fig. 3 as a function of the photon energy. The F channel remains small for all investigated PAH cations at all photon energies and is barely notable for the dicoronylene cation in Fig. 3. More specifically, Fig. 4 shows that the branching ratio (BR) for photoionization relative to photodissociation increases significantly with increasing N_C and reaches a minimal value of 0.98 for the largest studied cation. This trend of large PAHs differs from what was observed in our earlier study of medium-sized PAH cations for which a larger fraction of fragments was observed (Zhen et al. 2016). It is in line with the ionization BR of about 0.97 at 20 eV which was derived by Zhen et al. (2015) for the HBC cation, $C_{42}H_{18}^+$, by operating their home-made ion trap setup at the DESIRS beamline. The authors also reported a value of 0.70 ± 0.10 for the ionization BR of the ovalene cation at 20 eV, which can be compared to a value of 0.87 ± 0.02 in our experiment. This difference can be interpreted by the low mass resolution achieved in the former experiments which impacted both the isolation of the ^{12}C parent isotopomer before irradiation and the quantification of the abundance of $-H$ fragments in the photoproducts.

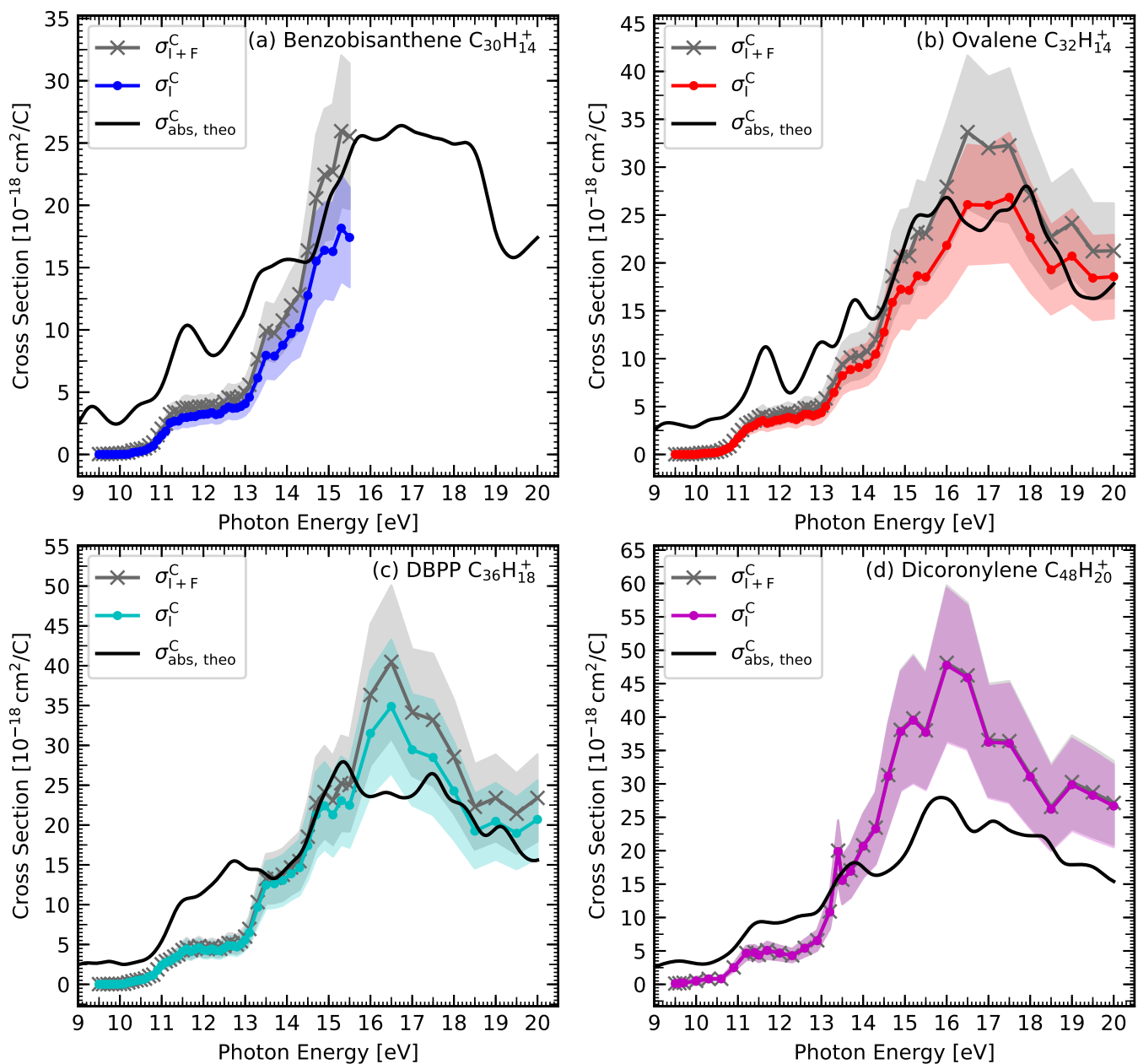


Fig. 5. Experimentally obtained total action and ionization cross sections per C atom, σ_{I+F}^C and σ_I^C , respectively, of the studied PAH cations as a function of photon energy, for (a) benzobisanthene, (b) ovalene, (c) DBPP, and (d) dicoronylene after absorption of a single VUV photon, compared to their by TD-DFT computed theoretical photoabsorption cross sections per C atom, $\sigma_{abs, theo}^C$.

Table 1. For doubly ionized PAH cations, $PAH^+ \rightarrow PAH^{2+}$, values of the theoretical adiabatic ionization potentials, IP^{2+} , and measured appearance energies, AE^{2+} . We also list here our recorded AE^{3+} as obtained from the ionization of $PAH^{2+} \rightarrow PAH^{3+}$.

Formula	IP^{2+} [eV]	AE^{2+} [eV]	AE^{3+} [eV]
$C_{30}H_{14}^+$	9.66 ^a	10.2 ± 0.1	14.1 ± 0.2
$C_{32}H_{14}^+$	9.82 ^a	10.0 ± 0.1	13.9 ± 0.2
$C_{36}H_{18}^+$	9.94 ^b	10.0 ± 0.1	13.9 ± 0.2
$C_{48}H_{20}^+$	8.84 ^a	9.1 ± 0.2	— ^c

Notes.

^(a) Taken from Mallocci et al. (2007a).

^(b) Calculated for this work according to Mallocci et al. (2004).

^(c) Trication peak out of mass range.

From Figure 3, we derived appearance energies for the formation of PAH^{2+} from PAH^+ , AE^{2+} . The values are listed in Table 1 and compared to the corresponding computed values for the adiabatic ionization potentials, IP^{2+} , which are extracted from the Theoretical Spectral Database of PAHs¹ (Mallocci et al. 2007a) or calculated at the same level of theory for the missing IP^{2+} of the DBPP cation according to Mallocci et al. (2007b). Experimental and theoretical values are found to be in good agreement considering the accuracy of ~ 0.3 eV for the calculated values. The trend of a slow decrease of IP^{2+} with $N_C \gtrsim 30$ carbon atoms reported by Mallocci et al. (2007b) is confirmed.

¹<http://astrochemistry.oa-cagliari.inaf.it/database/>

3.2. Photoionization cross sections

Experimental total action cross sections per carbon atom, σ_{I+F}^C , were obtained following the procedure described in Appendix C. The photoionization cross sections, σ_I^C , were then derived by using the branching ratio depicted in Fig. 4. The σ_{I+F}^C curves are expected to provide a lower value for the absolute photoabsorption cross sections, σ_{abs}^C (see Eq. (C.1)). Since σ_{abs}^C of the studied cations could not be extracted from the performed experiment and have not been reported so far in the literature, we compare these curves with the theoretical photoabsorption cross sections, $\sigma_{abs,theo}^C$, which have been computed using Time-Dependent Density Functional Theory (TD-DFT) in line with our previous work (Mallocci et al. 2004, 2007a) and as described in Appendix E. All obtained cross sections, experimental and theoretical, are displayed in Fig. 5 and compared to each other in the following at high (> 14 eV) and low (< 14 eV) energies.

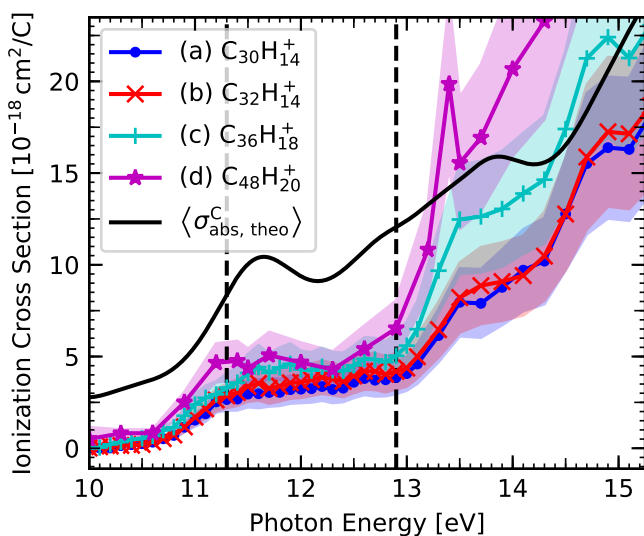


Fig. 6. Experimentally obtained photoionization cross sections per C atom, σ_I^C , of the studied PAH cations as a function of photon energy, (a) benzobisanthene, (b) ovalene, (c) DBPP, and (d) dicoronylene after absorption of a single VUV photon. The average of the calculated photoabsorption cross sections, $\langle \sigma_{abs,theo}^C \rangle$, is also presented. The dashed vertical lines mark the transitions between the different ionization regimes.

Above 14 eV, the cross sections are globally consistent (see Fig. 5). Still, the values of σ_{I+F}^C are found to be systematically larger than those of $\sigma_{abs,theo}^C$ around the peak at 17 eV. In addition, there is a trend of increasing σ_{I+F}^C at the peak with molecular size. The case of $C_{48}H_{20}^+$ has to be considered with caution though due to a less accurate calibration procedure (see Appendix C). On the contrary to the experimental cross sections, the values of $\sigma_{abs,theo}^C$ stay close to each other, which is expected from the proportionality of the photoabsorption cross sections with N_C . Still, it is not yet possible to access how precise the calculated cross sections are. The comparison with an experimental photoabsorption cross section at high energies (up to 30 eV) has been done so far only for neutral anthracene, $C_{14}H_{10}$ (Mallocci et al. 2004). It is interesting to mention that this comparison reveals an overall good agreement between the calculated and measured cross sections but with differences on the band positions and widths (in the theoretical spectra the band width is artificial). Also around the high energy peak observed at 18 eV,

the discrepancy appears similar to the one illustrated in Fig. 5 in the case of $C_{32}H_{14}^+$ and $C_{36}H_{18}^+$.

Below 14 eV, σ_{I+F}^C and σ_I^C are comparable since the fragmentation is negligible. A plateau is observed from 11.3 to 12.9 eV in these experimental cross sections whereas the theoretical photoabsorption cross section exhibits strong bands (see Fig. 5). There may be a number of possible reasons for the absence of strong, discrete absorption bands in σ_I^C . Ionization cross sections in PAHs may display considerable structure, corresponding both to autoionization resonances and to the opening of channels as energy increases, making accessible additional excited states in the resulting ion with one electron less, i.e., vertical transitions (see e.g. Bréchnac et al. 2014).

Theoretical considerations have shown that the orbital picture of ionization involving valence one-electron bands is severely contaminated by shake-up states, which involve two electrons, one promoted to an excited bound state and the other to the ionization continuum (Deleuze et al. 2001). The authors have shown that, in the case of π orbitals, this happens at energies as low as 8 eV for the first ionization (case of neutral PAHs). The density of excited states becomes quickly very large, blending in a quasi-continuum as energy increases (Deleuze et al. 2001). In addition, each electronic excited state may display vibrational structure on top of it, i.e., vibronic states, producing further structure on a smaller scale. One thus expects, qualitatively, a jump in the ionization cross section when a major channel becomes accessible via a valence one-electron transition, followed by a long plateau-like tail produced by all states coupling to it due to electronic correlation and vibronic coupling. On top of this, there may be electronic transitions to excited states of the parent molecule which have low coupling with the ionization continuum, therefore preferentially relaxing via radiationless transitions to lower electronic states. These do not show in the ionization cross sections. The theoretical method used here to compute the absolute, total cross section does not distinguish among different categories of electronic excitation, they are in principle all included together (except vibronic coupling) and cannot be distinguished in these calculations. More complex techniques involving many-body theory can compute the structure of the ionization cross sections with considerable accuracy (see e.g. Baiardi et al. 2017, for a recent review). However, their computational cost would be extremely high for the species considered here, and they are out of the scope of the present work, in which such detailed structure is not resolved anyway.

Below 14 eV, it is clear that there is a part of the photoabsorption cross section that does not lead to ionization, which we referred to as σ^- in Eq. (C.1). These correspond to excitations which involve fast relaxation via a strong coupling between electronic states and with nuclear states. This leads to vibrational excitation of the parent ion, which is expected to subsequently relax its energy by radiative cooling since no fragmentation is observed. At energies higher than 14 eV, evidence for such transitions, if they exist, is hindered by the precision of our experimental and theoretical data, as discussed above.

4. Astrophysical recipes

4.1. Charge state of astro-PAHs

A couple of modeling studies on the charge state of astro-PAHs have considered that these species could reach the dication and marginally the trication states (Bakes et al. 2001; Weingartner & Draine 2001b; Andrews et al. 2016). In Figure 7, we compiled the ionization potentials ($IP^{(Z+1)+}$) from neutral to cation

($Z = 0$), cation to dication ($Z = 1$), and dication to trication ($Z = 2$), which have been obtained from calculations (Mallocci et al. 2007b), as well as IP^{exp} and $\text{AE}^{(Z+1)+}$ from experiments (Clar et al. 1981; Hager & Wallace 1988; Zhen et al. 2016, and this work). We compared this data set with two analytical descriptions. Both start from a classical model of the energy it takes to remove one electronic charge from a small, conducting particle. Weingartner & Draine (2001a) considered conducting spheres, and added an empirical correction term to account for both quantum effects and PAH geometry, which is planar and not spherical. This additional term was determined to fit a set of data on first and second ionization potentials of PAHs and led to

$$\text{IP}_{WD}^{(Z+1)+} = W + \frac{1}{4\pi\epsilon_0} \left[\left(Z + \frac{1}{2} \right) \frac{e^2}{a} + (Z+2) \frac{e^2}{a} \frac{0.03 \text{ nm}}{a} \right] \frac{1 \text{ C}}{e}, \quad (1)$$

where Z is the ion charge, ϵ_0 is the vacuum permittivity in $\text{F} \cdot \text{nm}^{-1}$, e is the elementary charge in C, W is the work function of bulk graphite, $W = 4.4 \text{ eV}$, and a , the effective radius in nm, is proportional to N_C via the relation

$$a = \sqrt[3]{\frac{N_C}{468}}. \quad (2)$$

In Eq. (1), the first term in square brackets corresponds to the classical conducting sphere model, the second term is the empirical correction. For the data set reported in Fig. 7, we found that, instead of $W = 4.4 \text{ eV}$, a value of $W = 3.9 \text{ eV}$ fits the data better (black curves in Fig. 7). A satisfactory model is obtained for all Z values considered ($Z = 0, 1, 2$). The empirical formula which was adjusted for $Z = 0, 1$ appears also to be appropriate for $Z = 2$. The adjustment that we made on the W value corresponds to a vertical shift and somehow depends on the considered data set. For instance, we can see that our reported values for $\text{AE}^{(Z+1)+}$ are systematically slightly above the DFT values (cf. Table 1). Still, our derived value for W appears in line with the values of about 4.0 eV , which were calculated for similarly sized PAHs by Kvashnin et al. (2013).

The second formalism to describe $\text{IP}^{(Z+1)+}$ is given by Bakes & Tielens (1994) who considered a thin conducting disk instead of a sphere yielding

$$\text{IP}_{BT}^{(Z+1)+} = W + \left(Z + \frac{1}{2} \right) \frac{25.1}{\sqrt{N_C}}. \quad (3)$$

The gray curves in Fig. 7 have been obtained from Eq. (3) (Bakes & Tielens 1994) and using $W = 3.9 \text{ eV}$, as derived above. The curve provides a very satisfactory description of the data for $Z = 0$, but tends to increasingly fail for higher Z values. This trend was already noticed by Weingartner & Draine (2001a) and we can see that the discrepancy even increases for $Z = 2$. Tuning the value of W does not change the shape of the curves and this emphasizes the need to include quantum effects in the estimation of $\text{IP}^{(Z+1)+}$.

We can see from Fig. 7 and Eq. (1) that a fraction of the photons absorbed in H I regions can induce ionization of PAH cations. Taking the absorption and ionization cross sections shown in Fig. 5 and considering the radiation field of the prototypical NGC 7023 NW PDR (Joblin et al. 2018), we can estimate that typically one photon over three absorbed in the $[10 - 13.6] \text{ eV}$ range by PAH cations with $N_C = 30 - 36$ will lead to ionization. The fraction of ionizing events will increase

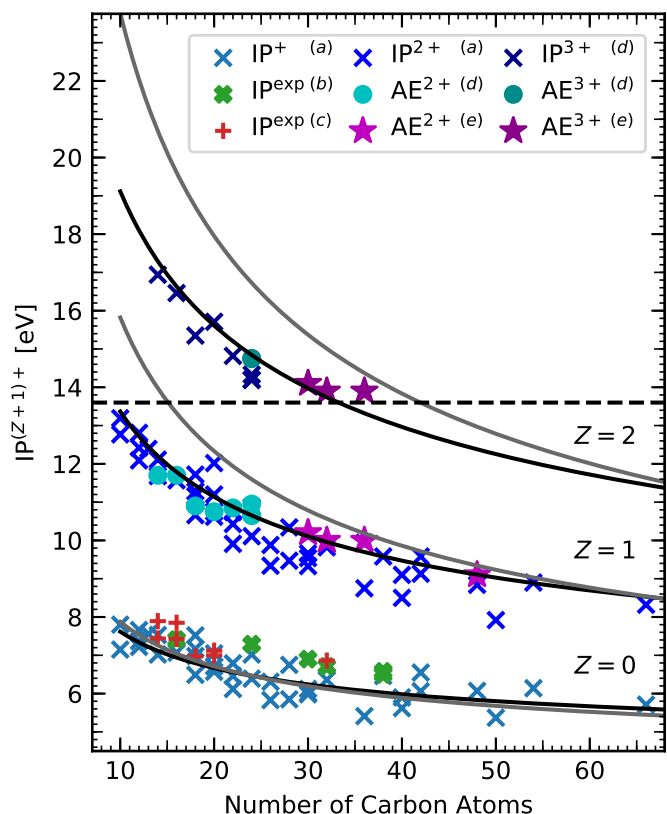


Fig. 7. Theoretically calculated ionization potentials, $\text{IP}^{(Z+1)+}$, experimentally obtained IP^{exp} and appearance energies, $\text{AE}^{(Z+1)+}$, as a function of number of carbon atoms. For $Z = 0, 1, 2$, the values refer to the transition from $\text{PAH}^{(Z)+} \rightarrow \text{PAH}^{(Z+1)+}$. The black and gray curves are the modeled $\text{IP}_{WD}^{(Z+1)+}$ from Weingartner & Draine (2001a) and $\text{IP}_{BT}^{(Z+1)+}$ from Bakes & Tielens (1994), respectively, as an estimate of the $\text{IP}^{(Z+1)+}$ evolution as a function of PAH size and charge, Z (see Eqs. (1) and (3) and text for details). The dashed horizontal line marks the 13.6 eV photon energy cut-off for H I regions.

Notes.

- (a) Taken from Mallocci et al. (2007b).
- (b) Taken from Clar et al. (1981).
- (c) Taken from Hager & Wallace (1988).
- (d) Taken from Zhen et al. (2016).
- (e) This work.

with increasing molecular size as the ionization potential shifts to lower energies. It reaches 0.5 for $\text{C}_{48}\text{H}_{20}^+$. We also note that the formation of C_{60}^{2+} will be more difficult to achieve than that of a PAH^{2+} of similar size, since the corresponding cations have relatively similar absorption cross sections but the value of AE^{2+} for C_{60}^+ is significantly higher, $(10.5 \pm 0.1) \text{ eV}$ (Douix et al. 2017), compared to 8.7 eV for a PAH^+ with $N_C = 60$ (see Fig. 7).

4.2. Photoionization yield

Photoionization yields of PAH cations were derived for all studied species by dividing the experimental photoionization cross section, σ_1^C , by the theoretical photoabsorption cross section, $\sigma_{\text{abs,theo}}^C$. In Sect. 3.2, we discussed the precision of both the experimental and theoretical cross sections. This can impact the photoionization yields. At energies below 14 eV , the presence of bands in $\sigma_{\text{abs,theo}}^C$, which are not present in σ_1^C , can induce spec-

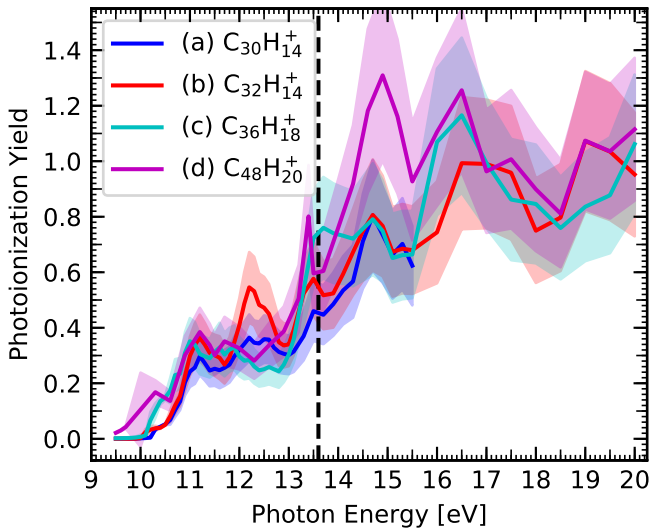


Fig. 8. Photoionization yields of the studied cations derived from the ionization and absorption cross sections (see Fig. 5) and scaled as explained in Sect. 4.2. The dashed vertical line marks the 13.6 eV photon energy cut-off for H I regions.

tral features in the photoionization yields (e.g. the 12 eV peak obtained for $C_{32}H_{14}^+$), which are as precise as the calculated spectrum. Still, Fig. 8 shows that the photoionization yields display comparable features for the studied molecules, with a rise starting at the ionization thresholds, AE^{2+} , the plateau in the 11.3 to 12.9 eV range followed by another rise to reach the maximum value. There is some uncertainty on this maximum value because of the unknown contribution from σ' (cf. Sect. 3.2). In the following, we made the hypothesis that the contribution of σ' at high energies (20 eV) is minor and that the photoionization yields are limited by the photoionization BR, which never reaches unity as shown in Fig. 4. The mean values at high energies of the photoionization yields were thus scaled using the ionization BR at 20 eV. The resulting curves are presented in Fig. 8.

Data on the photoionization yields of neutral PAHs have been previously derived from experimental studies performed by Verstraete et al. (1990) and Jochims et al. (1996). The latter authors have proposed a *rule of thumb* to facilitate the implementation of this yield into models. This consists of a linear function of the photon energy with dependence on the ionization potential. On the basis of our results, we propose to use a similar approach to describe the evolution of the photoionization yield of PAH cations with molecular size. The resulting function, Y_{ion}^+ , is based on the above described ionization regimes which occur in different energy ranges (values in eV) as

$$Y_{\text{ion}}^+[N_C](h\nu) = \begin{cases} 0 & h\nu < IP^{2+} \\ \frac{\alpha}{11.3 - IP^{2+}}(h\nu - IP^{2+}) & IP^{2+} \leq h\nu < 11.3 \\ \alpha & \text{for } 11.3 \leq h\nu < 12.9 \\ \frac{\beta(N_C) - \alpha}{2.1}(h\nu - 12.9) + \alpha & 12.9 \leq h\nu < 15.0 \\ \beta(N_C) & h\nu \geq 15.0, \end{cases} \quad (4)$$

where $\alpha = 0.3$ is the value of the plateau and β depends on N_C with

$$\beta(N_C) = \begin{cases} 0.59 + 8.1 \cdot 10^{-3} N_C & \text{for } 32 \leq N_C < 50 \\ 1 & N_C \geq 50. \end{cases} \quad (5)$$

The reported β values represent the values at 20 eV of the ionization BR (Fig. 4). They can be considered as maximum values

since they neglect a possible contribution of σ' to the photoabsorption cross section as discussed above. These values were found to increase linearly with size for the studied size range with the dependence given by Eq. (5). Extrapolation to larger sizes leads to a β value of 1 for $N_C \geq 50$. This trend differs from the case of neutral PAHs for which Jochims et al. (1996) concluded that $\beta = 1$ is independent of size, in agreement with previous measurements by Verstraete et al. (1990).

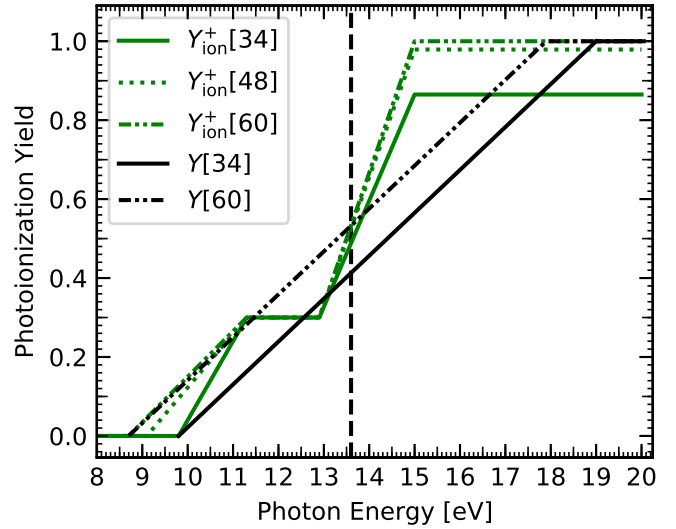


Fig. 9. Photoionization yields, $Y_{\text{ion}}^+[N_C](h\nu)$, calculated from Eqs. (4) and (5) for PAH cations with $N_C = 34, 48,$ and 60 atoms. For comparison, the photoionization yields, $Y[N_C](h\nu)$, for $N_C = 34$ and 60 , are displayed. They were estimated using the Eqs. (4) and (5) from Jochims et al. (1996) adapted for neutral PAHs but taking into account the shift of the ionization potential to IP^{2+} , which is relevant for cations. The dashed vertical line marks the 13.6 eV photon energy cut-off for H I regions.

Figure 9 displays examples of $Y_{\text{ion}}^+[N_C](h\nu)$ which were calculated from Eqs. (4) and (5), illustrating the variability of $Y_{\text{ion}}^+[N_C](h\nu)$ with molecular size. No significant variation of this yield is expected for PAH cations with $N_C \geq 50$. In their PAH evolution model, Andrews et al. (2016) have considered the yield of PAH cations based on the recipe given by Jochims et al. (1996) for neutrals but taking into account the appropriate photoionization potential for cations, i.e., values of IP^{2+} . To illustrate the impact that this approximation may have on the model results, we report in Fig. 9 these estimated yields and compare them with our recommended yields by integrating from IP^{2+} to 13.6 eV. We found that for the medium-sized PAHs, as represented by $N_C = 34$, our integrated yield is larger by 19% compared to the previously available one, whereas for large PAHs, as represented by $N_C = 60$, it is smaller by 14%. These simple estimates are however not conclusive and models have to be run to evaluate the impact on the ionization of the PAH population in specific environments.

5. Conclusion

We have studied the interaction of trapped PAH cations with VUV photons in the range of 9 to 20 eV, covering also photon energies present in H II regions and ionization fronts. Our experimental results provide a wealth of information on both ionization and fragmentation processes. The present article is focused on the detailed analysis of ionization, whereas fragmentation will

be the subject of a future, dedicated work. Our initial goal was to explore the properties of large species for N_C up to about 80 atoms. However we could only achieve measurements on molecular sizes from 30 to 48 carbon atoms due to the very low solubility of large PAHs. Still, studies in this range allow us to access the major trends in the ionization properties of PAH cations due to a molecular size increase. We found that

- (i) below 13.6 eV, the formation of a hot ion with subsequent (radiative) cooling is the major relaxation channel, followed by ionization whose yield reaches about 0.5 at 13.6 eV. From a molecular physics point of view, the yield comprises an interesting plateau at a value of 0.3 that extends over the energy range from 11.3 to 12.9 eV. This plateau reveals a spectral range in which there is a strong competition between electronic and nuclear states. It would be interesting to investigate the dynamics of the relaxation of excited electronic states in this range using fs pump-probe experiments (Marciniak et al. 2015).
- (ii) contrary to previous studies on neutrals, we could not observe that the photoionization yield reaches a value of 1 at high energies. At 20 eV, some dissociation is observed for all studied PAH cations, implying that the maximum of the yield cannot be larger than the branching ratio between ionization and dissociation, which increases with molecular size and reaches 0.98 for the largest studied ion, $C_{48}H_{20}^+$. In addition, we have not included a possible contribution in the photoabsorption events of the formation of a hot ion that would subsequently relax by radiative cooling in isolated conditions. This contribution would further lower the values of the photoionization yield. We have no explanation for the difference observed between neutrals and cations. Whether this is due to a change in their respective properties or the fact that experiments like ours using ion trapping are more sensitive to quantify this effect than experiments carried out on neutrals with different techniques, is out of our reach and would be interesting to further investigate.

Concerning astrophysical applications, we provide recipes to determine both the ionization potential and the photoionization yield of PAH cations as a function of their molecular size, which can be extended to larger sizes (typically $N_C = 100$). This yield can be combined with photoabsorption cross sections that are readily available from calculations using TD-DFT. All this molecular data can be used in models that describe the chemical evolution of PAHs in astrophysical environments. The range of photon energy we studied makes it possible to tackle the evolution of PAHs in extreme astronomical environments such as H II regions and ionization fronts. Observations of PAHs in these environments have been so far scarce due to their technical difficulty, but will become much more accessible thanks to the unique capabilities of the forthcoming James Webb Space Telescope (JWST).

As another example, the cavity around the star in NGC 7023 is expected to be an environment in which large PAH⁺ and PAH²⁺ are present (Andrews et al. 2016; Croiset et al. 2016). The presence of dications is expected to impact both the heating of the gas by photoelectric effect and the AIB emission. Some first IR action spectra of large PAH cations and dications have been recorded by Zhen et al. (2017, 2018). They provide encouraging results about large ionized PAHs being good candidates for carrying the AIBs. It is still not clear though if the spectral differences between cations and dications will be sufficient to differentiate both charge states in the observations. Still, we can predict that a detailed modeling approach combined with the wealth of

spectral and spatial information, which will be delivered soon by the JWST, will be able to highlight the charge evolution of the PAH population and its impact on the physics and chemistry of PDRs.

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References

- Andrews, H., Boersma, C., Werner, M. W., et al. 2015, *ApJ*, 807, 99
 Andrews, H., Candian, A., & Tielens, A. G. G. M. 2016, *A&A*, 595, A23
 Baiardi, A., Paoloni, L., Barone, V., Zakrzewski, V. G., & Ortiz, J. V. 2017, *J. Chem. Theory Comput.*, 13, 3120
 Bakes, E. L. O. & Tielens, A. G. G. M. 1994, *ASPC*, 58, 412
 Bakes, E. L. O. & Tielens, A. G. G. M. 1994, *ApJ*, 427, 822
 Bakes, E. L. O., Tielens, A. G. G. M., & Charles W. Bauschlicher, J. 2001, *ApJ*, 556, 501
 Bréchnignac, P., Garcia, G. A., Falvo, C., et al. 2014, *J. Chem. Phys.*, 141, 164325
 Cataldo, F., Ursini, O., Angelini, G., & Iglesias-Groth, S. 2011, *Fuller. Nanotub. Car. N.*, 19, 713
 Clar, E., Robertson, J. M., Schloegl, R., & Schmidt, W. 1981, *J. Am. Chem. Soc.*, 103, 1320
 Compiègne, M., Abergel, A., Verstraete, L., et al. 2007, *A&A*, 471, 205
 Croiset, B. A., Candian, A., Berné, O., & Tielens, A. G. G. M. 2016, *A&A*, 590, A26
 Deleuze, M. S., Trofimov, A. B., & Cederbaum, L. S. 2001, *J. Chem. Phys.*, 115, 5859
 Douix, S., Duflo, D., Cubaynes, D., Bizau, J.-M., & Giuliani, A. 2017, *J. Phys. Chem. Lett.*, 8, 7
 Giuliani, A., Giorgetta, J.-L., Ricaud, J.-P., et al. 2012, *Nucl. Instrum. Methods Phys. Res. B*, 279, 114
 Hager, J. W. & Wallace, S. C. 1988, *Anal. Chem.*, 60, 5
 Joblin, C., Bron, E., Pinto, C., et al. 2018, *A&A*, 615, A129
 Jochims, H. W., Baumgaertel, H., & Leach, S. 1996, *A&A*, 314, 1003
 Kvashnin, D. G., Sorokin, P. B., Brüning, J. W., & Chernozatonskii, L. A. 2013, *Appl. Phys. Lett.*, 102, 183112
 Le Page, V., Snow, T. P., & Bierbaum, V. M. 2003, *ApJ*, 584, 316
 Mallocci, G., Joblin, C., & Mulas, G. 2007a, *Chem. Phys.*, 332, 353
 Mallocci, G., Joblin, C., & Mulas, G. 2007b, *A&A*, 462, 627
 Mallocci, G., Mulas, G., & Joblin, C. 2004, *A&A*, 426, 105
 Marciniak, A., Despré, V., Barillot, T., et al. 2015, *Nat. Commun.*, 6, 1
 Matsuzawa, N. N., Ishitani, A., Dixon, D. A., & Uda, T. 2001, *J. Phys. Chem. A*, 105, 4953
 Milosavljević, A. R., Nicolas, C., Gil, J.-F., et al. 2012, *J. Synchrotron Radiat.*, 19, 174
 Montillaud, J., Joblin, C., & Toublanc, D. 2013, *A&A*, 552, A15
 Nahon, L., de Oliveira, N., Garcia, G. A., et al. 2012, *J. Synchrotron Radiat.*, 19, 508
 Tancogne-Dejean, N., Oliveira, M. J. T., Andrade, X., et al. 2020, *J. Chem. Phys.*, 152, 124119
 Tielens, A. G. G. M. 2005, *The Physics and Chemistry of the Interstellar Medium* (Cambridge University Press)
 Verstraete, L., Leger, A., D'Hendecourt, L., Defourneau, D., & Dutuit, O. 1990, *A&A*, 237, 436
 Vicente, S., Berné, O., Tielens, A. G. G. M., et al. 2013, *ApJ*, 765, L38
 Visser, R., Geers, V. C., Dullemond, C. P., et al. 2007, *A&A*, 466, 229
 Weingartner, J. C. & Draine, B. T. 2001a, *ApJ*, 548, 296
 Weingartner, J. C. & Draine, B. T. 2001b, *ApJS*, 134, 263
 Yabana, K. & Bertsch, G. F. 1996, *Phys. Rev. B*, 54, 4484
 Zhen, J., Candian, A., Castellanos, P., et al. 2018, *ApJ*, 854, 27
 Zhen, J., Castellanos, P., Bouwman, J., Linnartz, H., & Tielens, A. G. G. M. 2017, *ApJ*, 836, 28
 Zhen, J., Castellanos, P., Paardekooper, D. M., et al. 2015, *ApJ*, 804, L7
 Zhen, J., Rodriguez Castillo, S., Joblin, C., et al. 2016, *ApJ*, 822, 113

Appendix A: Synthesis of dicoronylene, C₄₈H₂₀

Dicoronylene was synthesized following the procedure reported by Cataldo et al. (2011). A mixture of coronene (0.060 g, 0.199 mmol), AlCl₃ (4.20 g, 31.2 mmol), NaCl (0.817 g, 14.0 mmol), and CuCl (0.043 g, 0.440 mmol) was stirred at 195 °C for 2 h. Then, aqueous solution of HCl (10 %, 20 ml) was slowly added resulting in the precipitation of a red solid. This solid was filtrated and washed with aqueous solution of HCl (10 %, 20 ml), hot water (20 ml) and acetone (3 × 20 ml). The resulting solid was dried to obtain dicoronylene as a red solid (54 mg, 40 %).

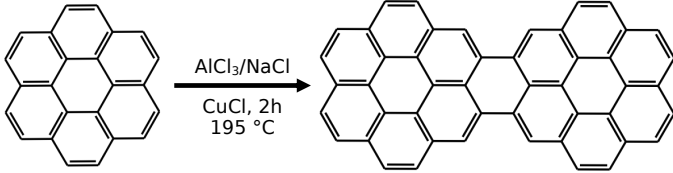


Fig. A.1. Schematic of the synthesis of dicoronylene, C₄₈H₂₀, from coronene, C₂₄H₁₂.

Appendix B: Action spectra scaling procedure

From the recorded mass spectra, peak intensities of parent ions and their photoproducts can be deduced. The secondary ions produced upon VUV irradiation consist of the dication with peak intensity S_I , and the photofragments with summed peak intensity S_F , including the H and 2H/H₂ loss channels. Due to detector characteristics, doubly ionized molecules are detected more efficiently than singly ionized molecules. Therefore, the peak intensities derived from the mass spectra have to be scaled by the detector gain efficiency, ε , to retrieve values that scale with abundances. Thermo Scientific™ provides $\varepsilon_+ = 0.29$ for parent ions and fragments, $\varepsilon_{2+} = 0.42$ for dications and a value of 0.54 for trications. There is some gain change with mass but this is a minor correction for the range of studied masses.

The total number of ions in the trap, P_0 (in uncalibrated values), can then be calculated with

$$P_0 = \frac{P_t}{\varepsilon_+} + \frac{S_F}{\varepsilon_+} + \frac{S_I}{\varepsilon_{2+}}, \quad (\text{B.1})$$

where $\frac{P_t}{\varepsilon_+}$ is the number of parent ions after irradiation time, t . Building the action spectra requires to derive normalized photoproduct intensities for S_I and S_F , which can be obtained by first dividing them by P_0 and then correcting for the variation of the photon flux, $\Phi(\nu)$. Indeed the latter evolved in energy due to spectral shape variations and changes made in the irradiation time, t , and monochromator exit slit width, s , so that the total photoproduct intensity remains smaller than *ca.* 12 % of P_t , as seen in the mass spectra. This led to

$$S_{I,F}^{\text{norm}}(\nu) = \frac{S_{I,F}(\nu)}{\Phi_{\text{norm}}(\nu) t(\nu) s(\nu)}, \quad (\text{B.2})$$

where $\Phi_{\text{norm}}(\nu)$ is normalized to be 1 at its maximum at 9.5 eV. Note that the obtained intensities are in arbitrary units (see Fig. 3) and not in percentage of the total number of ions because of the scaling by the relative photon flux.

Appendix C: Photoproduct cross sections scaling procedure

The absorption of a photon by the parent ion leads to different relaxation channels. Considering monochromatic radiation, the photoabsorption cross section, σ_{abs} , can therefore be decomposed into the sum of the cross sections for each relaxation channel as

$$\sigma_{\text{abs}} = \sigma_{I+F} + \sigma', \quad (\text{C.1})$$

where σ_{I+F} is the cross section leading to the production of the secondary products with intensities S_I and S_F (see Appendix B), and σ' is the cross section for the creation of a hot ion that will relax its internal energy by radiative cooling and/or collisions with buffer gas (He) in our experiment. These processes cannot be traced in our experiment and only σ_{I+F} can be estimated following

$$\sigma_{I+F} = \frac{\gamma}{\Phi t} \ln \left(\frac{\varepsilon_+ P_0}{P_t} \right), \quad (\text{C.2})$$

with the photon flux, Φ , in photons cm⁻²s⁻¹, a form factor, γ , describing the overlap of the photon beam and the ion cloud, the total number of ions, P_0 , and the number of parent ions, $\frac{P_t}{\varepsilon_+}$, after irradiation time, t . Plugging P_0 from Eq. (B.1) in Eq. (C.2), we get

$$\sigma_{I+F} = \frac{\gamma}{\Phi t} \ln \left(1 + \frac{S_F}{P_t} + \frac{\varepsilon_+ S_I}{\varepsilon_{2+} P_t} \right). \quad (\text{C.3})$$

In order to determine cross sections in absolute units, the photon flux, Φ , and the form factor, γ , have to be well-known. Douix et al. (2017) managed to record the absolute photoionization cross section, $\sigma_1^{C_{60}^+}$, for the buckminsterfullerene cation, C₆₀⁺, by carefully measuring these parameters and applying Eq. (C.3), where the term $\frac{S_F}{P_t}$ was zero due to the non-dissociation of C₆₀²⁺. In our experiment, we trapped C₆₀⁺ and recorded its dication peak under the same irradiation conditions used for our studied PAHs except dicoronylene as explained below. We could therefore derive the value of $\frac{\gamma}{\Phi t}$ and use this factor to obtain experimental values for σ_{I+F} from the photoproduct evolutions of our PAHs according to Eq. (C.3). This used scaling procedure yields reasonable cross section values only above the AE²⁺ of C₆₀⁺. We bypassed this limitation by scaling the action spectra (see Sect. 3.1 and Appendix B) to the cross sections and replacing the values below the AE²⁺ of C₆₀⁺ with the values from the scaled action spectra. Finally, to simplify the comparison between species, we divide σ_{I+F} of each PAH cation by its respective N_C , yielding σ_{I+F}^C .

We note that in this calibration procedure, dicoronylene required a specific treatment. Indeed this ion was studied in different conditions since both the syringe flow rate and the monochromator exit slit width, s , were increased in order to get a sufficient signal. In order to correct at best for these changes, we applied corrections to the $\frac{\gamma}{\Phi t}$ factor in Eq. (C.2) by assuming that not only the photon flux but also the beam overlap with the ion cloud scales linearly with s , the later is likely disputable but this is the best we could do.

²The absolute photoionization cross section of C₆₀⁺, $\sigma_1^{C_{60}^+}$, can be accessed at 10.5281/zenodo.1001072

Appendix D: Error estimation

Depending on the acquisition time, a few hundred, N , mass spectra, x , are recorded for each photon energy step. Averaging these mass spectra for each photon energy yields one mean mass spectrum, \bar{x} , per photon energy with an absolute standard error of $\Delta\bar{x} = \frac{\sigma}{\sqrt{N}}$, where σ is the standard deviation. The error bars, Δf , in Figs. 3 and 4 result from error propagation for a function, $f \rightarrow f(x_1, x_2, \dots, x_n)$, according to

$$\Delta f = \sqrt{\sum_{i=1}^n \left(\frac{\partial f}{\partial x_i} \Delta x_i \right)^2}. \quad (\text{D.1})$$

When determining the cross sections depicted in Fig. 5 following the procedure presented in Appendix C as well as the obtained photoionization yields shown in Figs. 8 and 9, the error on the cross section of C_{60}^+ , which was found to be below 21 % by Douix et al. (2017), is used to propagate the errors.

Appendix E: Theoretical cross section calculations

Theoretical photoabsorption cross sections were obtained using Time-Dependent Density Functional Theory (TD-DFT), with the real-time, real-space method of Yabana & Bertsch (1996), as implemented in the OCTOPUS computer code (Tancogne-Dejean et al. 2020). Not all species in this work were considered in our previous work (Mallocci et al. 2007a). In addition, since we here need to use theoretical spectra up to relatively high energies (~ 20 eV), we also took the chance to verify at these energies the convergence of the calculations with respect to the simulation box size and grid spacing, the real-space equivalent of more conventional basis-set convergence for e.g. Gaussian-based DFT. It is known (Matsuzawa et al. 2001), and mentioned in the OCTOPUS documentation that, as a rule of thumb, larger simulation boxes are needed to obtain converged values for the energies of higher lying electronic excited states, and for the intensities of the transitions involving them. Moreover, when using the real-time, real-space method implemented in OCTOPUS, the photoabsorption spectrum includes both transitions to discrete (bound) electronic states and to the continuum of unbound states. The former tend to converge to well-defined energies in the limit of infinite box size; the latter, instead, are artificially quantized by the boundary conditions of a finite simulation box, whose individual energies keep on changing with changing box size (and shape).

In this context, we have started a computational effort to test convergence as a function of box sizes. It led us to conclude that the spectra in the database of Mallocci et al. (2007a) are adequately converged for excitation energies up to ~ 13 eV, meaning they are suitable for astronomical modeling purposes involving photons up to the Lyman limit. We instead had to consider substantially larger simulation box sizes, and denser grid spacing, to achieve convergence up to ~ 20 eV, as required for this work. For our purposes, we found that a “minimal” simulation box can be achieved by considering the union of 8 Å radius spheres centered on each atom of the molecule, or by a single larger sphere with an equivalent volume. In addition, a grid spacing of 0.18 Å was needed to achieve convergence. All the calculations were performed on the OLYMPE supercomputer at the mesocentre CALcul en MIDI Pyrénées (CALMIP). We retained from the set of simulations two different converged spectra for each species and averaged them together in order to smooth out the contribution of free electron states and retain discrete ones. Finally, we divided

the $\sigma_{\text{abs,theo}}$ of each PAH cation by its respective N_{C} , yielding $\sigma_{\text{abs,theo}}^{\text{C}}$.