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MKT J170456.2–482100: the first transient discovered by MeerKAT

L. N. Driessen¹,^{*} I. McDonald¹, D. A. H. Buckley², M. Caleb¹, E. J. Kotze^{2,3},
S. B. Potter², K. M. Rajwade¹, A. Rowlinson^{4,5}, B. W. Stappers¹, E. Tremou⁶,
P. A. Woudt⁷, R. P. Fender^{7,8}, R. Armstrong^{7,9}, P. Groot^{2,7,10}, I. Heywood^{8,11},
A. Horesh¹², A. J. van der Horst^{13,14}, E. Koerding¹⁰, V. A. McBride^{2,15,16},
J. C. A. Miller-Jones¹⁷, K. P. Mooley^{18,19,20} and R. A. M. J. Wijers⁴

Affiliations are listed at the end of the paper

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ABSTRACT

We report the discovery of the first transient with MeerKAT, MKT J170456.2–482100, discovered in ThunderKAT images of the low-mass X-ray binary GX339–4. MKT J170456.2–482100 is variable in the radio, reaching a maximum flux density of 0.71 ± 0.11 mJy on 2019 October 12, and is undetected in 15 out of 48 ThunderKAT epochs. MKT J170456.2–482100 is coincident with the chromospherically active K-type sub-giant TYC 8332-2529-1, and ~ 18 yr of archival optical photometry of the star shows that it varies with a period of 21.25 ± 0.04 d. The shape and phase of the optical light curve changes over time, and we detect both X-ray and UV emission at the position of MKT J170456.2–482100, which may indicate that TYC 8332-2529-1 has large star spots. Spectroscopic analysis shows that TYC 8332-2529-1 is in a binary, and has a line-of-sight radial velocity amplitude of 43 km s^{-1} . We also observe a spectral feature in antiphase with the K-type sub-giant, with a line-of-sight radial velocity amplitude of $\sim 12 \pm 10 \text{ km s}^{-1}$, whose origins cannot currently be explained. Further observations and investigation are required to determine the nature of the MKT J170456.2–482100 system.

Key words: stars: activity – binaries: spectroscopic – stars: flare – stars: peculiar.

1 INTRODUCTION

The radio sky contains many variable and transient sources, often found in follow-up observations of transients detected at other wavelengths such as optical, gamma-ray, and X-ray (e.g. Sood & Campbell-Wilson 1994; Zauderer et al. 2011; Chandra & Frail 2012; Horesh et al. 2013; Fong et al. 2015; Marsh et al. 2016; Hallinan et al. 2017; Bright et al. 2019). Blind searches for radio transients using interferometers present many challenges, particularly modest field of view (FoV) and limited observing cadence (e.g. Murphy et al. 2013; Mooley et al. 2016, 2018). With current wide FoV ($\gtrsim 1 \text{ deg}^2$) instruments such as MeerKAT (Camilo et al. 2018), the Australian Square Kilometer Array Pathfinder (ASKAP; Johnston et al. 2008; Schinckel et al. 2012), APERTIF (Maan & van Leeuwen 2017), the LOw Frequency Array (LOFAR; van Haarlem et al. 2013), and the Murchison Wide Field Array (MWA; Tingay et al. 2012), surveying large areas of sky with various cadences and improved sensitivity is now possible. These new instruments could result in the discovery of tens to hundreds of transients (e.g. O’Brien et al. 2015).

Radio transients are commonly divided into two categories: coherent and incoherent (e.g. Pietka, Fender & Keane 2015); and both types of transient are investigated in the time domain with high-time resolution (milliseconds or less), and in image plane observations with a range of integration time-scales. In this publication we will focus on image plane searches. Current image plane transient searches include the Amsterdam-ASTRON Radio Transients Facility and Analysis Centre (AARTFAAC; Prasad et al. 2016; Kuiack et al. 2019), and the ASKAP Survey for Variables and Slow Transients (VAST; Murphy et al. 2013). Large surveys such as the Very Large Array (VLA) Sky Survey (VLASS; Lacy et al. 2019) are also being used to search for transients (Hallinan et al. 2013). It was originally theorized that image plane, low-frequency transient searches would detect many transient radio sources, but to date only one transient each has been found with LOFAR (Carbone et al. 2016; Stewart et al. 2016), the Long Wavelength Array (LWA; Varghese et al. 2019) and the MWA (Murphy et al. 2017), and no transients have been found with the VLA Low Band Ionospheric and Transient Experiment (VLITE; Polisensky et al. 2016). The rate of low-frequency Galactic transients may be higher, as inferred from the Galactic Center Radio Transients detected by VLA and Giant Metrewave Radio Telescope (GMRT; e.g. Hyman et al. 2005, 2009;

* E-mail: laura@driessen.net.au

Roy et al. 2010). At higher frequencies, a VLA search for transients at 5 GHz only found a single transient candidate (Ofek et al. 2011), while the Caltech-NRAO Stripe 82 Survey Pilot (CNSS; Mooley et al. 2016) at 3 GHz detected two transients, and several transients have been found in the full CNSS survey (Mooley et al. in preparation). Bower et al. (2007) searched 944 epochs over 22 yr of VLA 5 and 8 GHz observations and found 10 new transients; however, more than half of these were found to be artefacts by Frail et al. (2012). More recently, nine potential variable sources were detected using ASKAP (Bhandari et al. 2018). This means that only a few radio transients have been discovered in blind image plane transient surveys, despite expectations for many new discoveries. These results¹ have highlighted the importance of wide-field, sensitive searches, at suitable frequencies, for maximizing the yield of radio transients (see also Bannister et al. 2011; Thyagarajan et al. 2011).

One type of radio transient expected to be found in image plane transient searches is flares from stars and stellar systems (see e.g. Osten 2008, for a summary). Radio flare stars are usually M-type dwarf stars that emit coherent radio bursts on time-scales of minutes to hours. Recently, Villadsen & Hallinan (2019) detected 22 coherent radio bursts from M dwarfs using the VLA at 300 MHz and 1–6 GHz, and Zic et al. (2019) detected several pulses from the M-dwarf UV Ceti with ASKAP. As well as flare stars, binary systems such as RS Canum Venaticorum (RS CVn), cataclysmic variables (CVs), and symbiotic binaries are known to flare in the radio. RS CVn are binary systems consisting of a late-type giant or sub-giant star with a late-type main-sequence star companion (e.g. Hall 1976; Craig et al. 1997; García-Sánchez, Paredes & Ribó 2003). RS CVn are chromospherically active and the giant or sub-giant rotates quasi-synchronously with the orbital period. Periods for RS CVn are typically 1–20 d, and radio flares on RS CVn systems have been observed to last up to a few days (e.g. Walter et al. 1987; Triglio, Umana & Migenes 1993). RS CVn emit in both the radio and X-ray while in quiescence (e.g. Gunn 1996). CVs are binary systems with a white dwarf primary accreting matter from Roche lobe overflow of the secondary star, usually a main-sequence star (e.g. Coppejans et al. 2016). Dwarf novae from magnetic CVs have been observed to exhibit radio outbursts that can last for weeks, with rapid radio flaring on time-scales less than an hour (Mooley et al. 2017). Symbiotic binaries also have a white dwarf primary, but the companion is a red giant star and the orbit is wide. In these systems, mass is accreted on to the white dwarf via stellar winds (O’Brien et al. 2015). Radio variability over time-scales of a few hours have been detected on symbiotic binaries, for example RX Puppis (Seaquist 1977). AR Scorpii is another stellar binary that is observed in the radio, it consists of a white dwarf primary and an M-dwarf companion. AR Scorpii has an orbital period of 3.56 h, is observed to pulsate in the optical, radio and X-ray with a period of 1.97 min (Marsh et al. 2016; Takata et al. 2018), and is highly polarized (e.g. Buckley et al. 2017). Radio flaring stellar systems vary on a variety of time-scales, making it difficult to detect when these systems are in outburst (e.g. Osten & Bastian 2008). This means that wide-field monitoring is an important method for discovering and investigating these sources in the radio.

MeerKAT is the (more) Karoo Array Telescope (Camilo et al. 2018) consisting of 64, 13.96 m dishes in South Africa. MeerKAT has an FoV of over a square degree at 1.4 GHz, which makes it an excellent instrument for searching for radio transients, such as

stellar flares. ThunderKAT is the MeerKAT Large Survey Project (LSP) investigating and searching for transients in the image plane (Fender et al. 2017). ThunderKAT is directly observing transient sources such as X-ray binaries, CVs, and gamma-ray bursts, and is commensally searching for radio transients. ThunderKAT has committed to observing the low-mass X-ray binary GX339–4 every week for 5 yr. With an FoV of over a square degree, weekly observations, and hundreds of sources in the field, this is an excellent opportunity to search for transient and variable radio sources that vary on many different time-scales. In this paper we will present the serendipitous discovery of MKT J170456.2–482100, the first radio transient discovered commensally with MeerKAT, in the GX339–4 field. In Section 2 we will present the radio observations of the source and the method of detection. In Section 3 we will identify the optical counterpart to MKT J170456.2–482100. In Sections 4, 5, 6, 7, and 8 we will present the optical photometry, optical spectroscopy, radio pulsation searches, UV photometry, and X-ray photometry of the source, respectively. In Sections 9 and 10 we will discuss our findings and conclude.

2 MEERKAT RADIO OBSERVATIONS

ThunderKAT (Fender et al. 2017) first observed the GX339–4 (Tremou et al. in preparation) field with 16 dishes on 2017 November 11 during commissioning. It then observed the field with all 64 dishes for the first time on 2018 April 14, and it began weekly monitoring in 2018 September. We present data from 48 epochs, 46 epochs from weekly monitoring plus the 2017 November 11 and 2018 April 14 epochs. The GX339–4 field is observed using the L-band (900–1670 MHz) receiver in full polarization mode, which has a bandwidth of 856 MHz, a central frequency of 1284 MHz, and 4096 frequency channels. The observations are typically 10–15 min in duration, with a minimum integration time of 8 s. The phase calibrator (1722–554) is observed for ~ 2 min before or after observing the field, and the band-pass and flux calibrator (1934–638) is observed for 5 min at the beginning of the observing block.

Each observation is first flagged using AOFFLAGGER² (Offringa 2010) and calibrated by following standard procedures (a priori phase correction, antenna delays and band-pass corrections) using the Common Astronomy Software Application³ (CASA; McMullin et al. 2007). As the data volume is large, the data are binned to reduce the number of channels from 4096 to 512. The full ThunderKAT pipeline will be presented in a future publication.

The data are imaged using the new wide-band, wide-field imager, DDFACET (Tasse et al. 2018). DDFACET is based on a co-planar faceting scheme, and it takes generic direction-dependent effects into account. This is important as MeerKAT has a very wide (~ 1 deg²) FoV. The imaging was performed by deconvolving over four frequency subbands using the SSDCLEAN deconvolution algorithm and a Briggs weighting scheme (robust = -0.7). To correct for considerable artefacts from bright sources, self-calibration was performed using the KILLMS⁴ software that accompanies DDFACET. The image quality was optimized by using the Complex Half-Jacobian Optimisation for N-directional ESTimation (COHJONES; Smirnov & Tasse 2015) algorithm to correct for direction-dependent effects. The COHJONES algorithm solves for scalar Jones matrices in a user-defined number of directions; three directions were used

¹A comprehensive list of blind radio transient searches can be found at <http://www.tauceti.caltech.edu/kunal/radio-transient-surveys/index.html>.

²<https://sourceforge.net/projects/aoflagger/>

³<https://casa.nrao.edu/>

⁴<https://github.com/saopicc/killms>

for the GX339–4 field. For the final images of the field the median synthesized beam in L band is $5 \text{ arcsec} \times 4 \text{ arcsec}$. The images typically have an overall root-mean-square (RMS) noise of $\sim 31.7 \mu\text{Jy beam}^{-1}$. We have found that there is a flux-dependent underlying systematic flux density fluctuation of up to the 10 percent in the light curves of sources in this field, likely caused by flux density variation in the secondary calibrator source. This does not affect the results of the analysis in this work; however, any small-scale variability visible in the MeerKAT light curves presented here should not be over-interpreted.

2.1 Transient detection pipeline

In 2018 December the images from the regular monitoring of GX339–4 were used to trial the LOFAR TRANSIENT PIPELINE (TRAP, Release 4.0; Swinbank et al. 2015) for detecting variable and transient sources in MeerKAT observations. TRAP automatically processes a time series of images by finding sources, determining source associations, and constructing light curves. The light curves are then used to calculate two variability parameters: η and V . The η parameter is the reduced χ^2 value for a fit to a stable source, higher η values show increased variation from a constant flux density while lower η values indicate that the flux density is stable. The V parameter is the co-efficient of variability or the modulation parameter, which is defined as the ratio between the standard deviation in the flux density and the mean flux density value. A higher V value indicates larger fractional variation in flux density values. For further information on the processing steps and the calculation of the variability parameters refer to Swinbank et al. (2015).

The default TRAP settings were used to analyse the GX339–4 field.⁵ A source detection threshold of 8σ was used, and the source finder was forced to use a Gaussian shape consistent with the synthesized beam of the images to search for sources. This was done to prevent the source finder from detecting extended emission, as variable sources are typically expected to be point sources. The first trial use of TRAP on the GX339–4 field was highly successful, the distributions of the variability parameters η and V are shown in Fig. 1. An outlier in both η and V was found to coincide with the known pulsar PSR J1703–4851, labelled in Fig. 1, at Right Ascension and Declination of $17^{\text{h}}03^{\text{m}}54^{\text{s}}.53(2)$ and $-48^{\text{d}}52^{\text{m}}01^{\text{s}}.0(5)$, respectively (J2000; Jankowski et al. 2019). This is a known mode-changing pulsar (Wang, Manchester & Johnston 2007), and this mode changing causes the flux density integrated over ~ 10 min to vary from epoch to epoch. An interesting source, MKT J170456.2–482100, was also identified as an outlier to the distribution of the V variability parameter. It was confirmed as a variable source by visual inspection of the light curve and images. At the time of discovery the variability parameters for MKT J170456.2–482100 were $\eta = 3.7$ and $V = 0.62$, while the point sources near MKT J170456.2–482100 are consistent with stable sources with $0.15 \leq \eta \leq 0.5$ and $0.04 \leq V \leq 0.09$. We note that there are other outliers in Fig. 1, these will be investigated and presented in a future publication.

Following the successful identification of MKT J170456.2–482100 using TRAP, the source was monitored in each weekly image of the GX339–4 field. The light curve of MKT J170456.2–482100, extracted using the TRAP default parameters, for all 48 epochs is shown in Fig. 2. Examples of

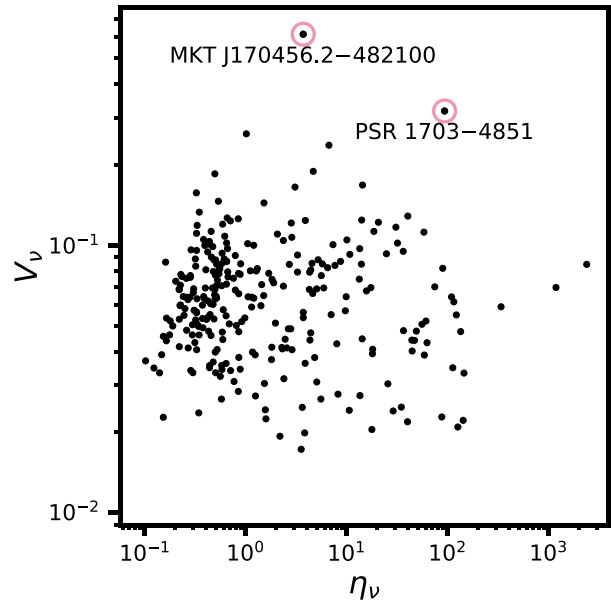


Figure 1. The variability parameters for all of the sources detected in the field of GX339–4. PSR J1703–4851, and MKT J170456.2–482100 (labelled) are clear outliers from the population of sources.

MeerKAT images of the source are shown in Fig. 3. To produce this light curve we forced TRAP to take measurements at the position of MKT J170456.2–482100 in every epoch, which means that TRAP measured values even when the source was not detected, using a detection threshold of 2σ above the noise in the image local to the source. In Fig. 2 we plot these non-detections (which were confirmed by eye) as 1σ upper-limits where σ is the RMS of the noise close to MKT J170456.2–482100. The uncertainties on the flux density are the 1σ uncertainties calculated by TRAP. The position of MKT J170456.2–482100 places it at approximately the half power point of the MeerKAT primary beam when GX339–4 is at the phase centre. On 2019 May 18 we observed GX339–4 as usual, and also took an observation with MKT J170456.2–482100 at the phase centre. Both observations were processed and calibrated in the same way. We then measured the flux density of the closest constant source to MKT J170456.2–482100 in both images, for ease we will call this nearby source 170 500–482 103. We determined that the flux density of 170 500–482 103 when MKT J170456.2–482100 is at the phase centre of the observation is 2.00 ± 0.20 times the flux density of 170 500–482 103 when GX339–4 is at the phase centre. This indicates that 170 500–482 103 and MKT J170456.2–482100 are at approximately the 50 percent point of the primary beam. Therefore, the flux density of MKT J170456.2–482100 in this manuscript is the measured flux density multiplied by 2.00 ± 0.20 to correct for the primary beam effect. Table A1 in Appendix A includes the flux density measurements of the peak flux density of the source as measured by TRAP, as well as the RMS noise measured near MKT J170456.2–482100.

We investigated shorter time-scale variability by dividing each epoch of observations into 160 s images and running these shorter integration images through TRAP. There is no variability in the 160 s integration images above one standard deviation from the mean of the light curve.

⁵For further details about the TRAP capabilities, refer to the TRAP documentation at <https://tkp.readthedocs.io/en/latest/> (TraP contributors 2014).

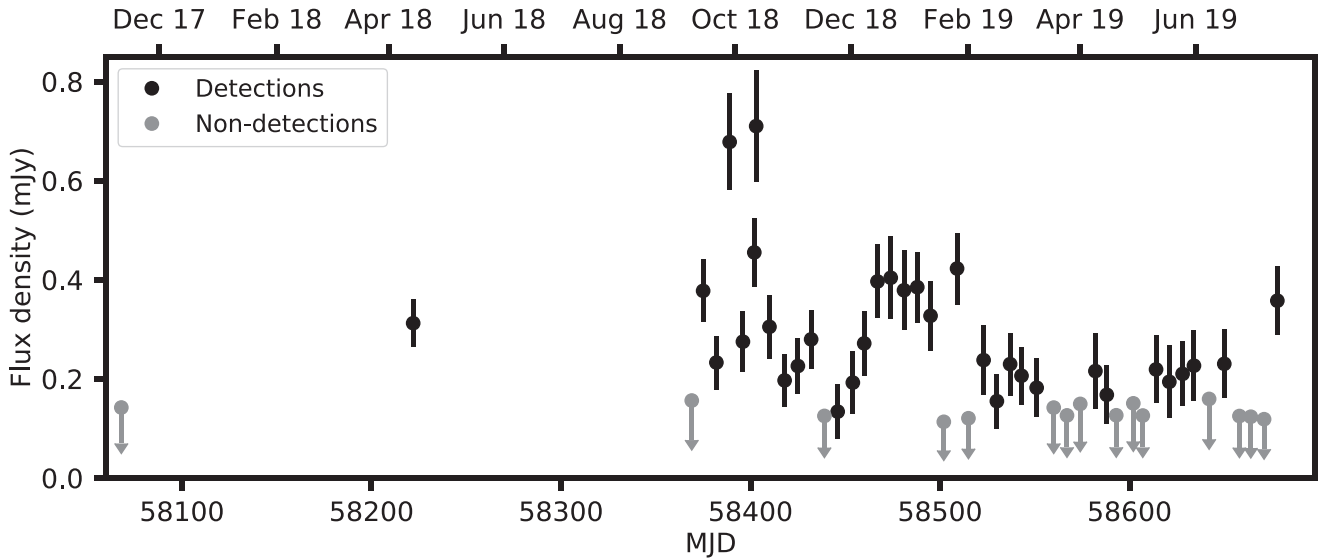


Figure 2. Radio light curve of MKT J170456.2–482100 from the ThunderKAT observations described in Section 2. We used a detection threshold of 2σ , where σ is the RMS noise close to the source. The upper limits are non-detections showing the 1σ RMS noise of the image local to the source.

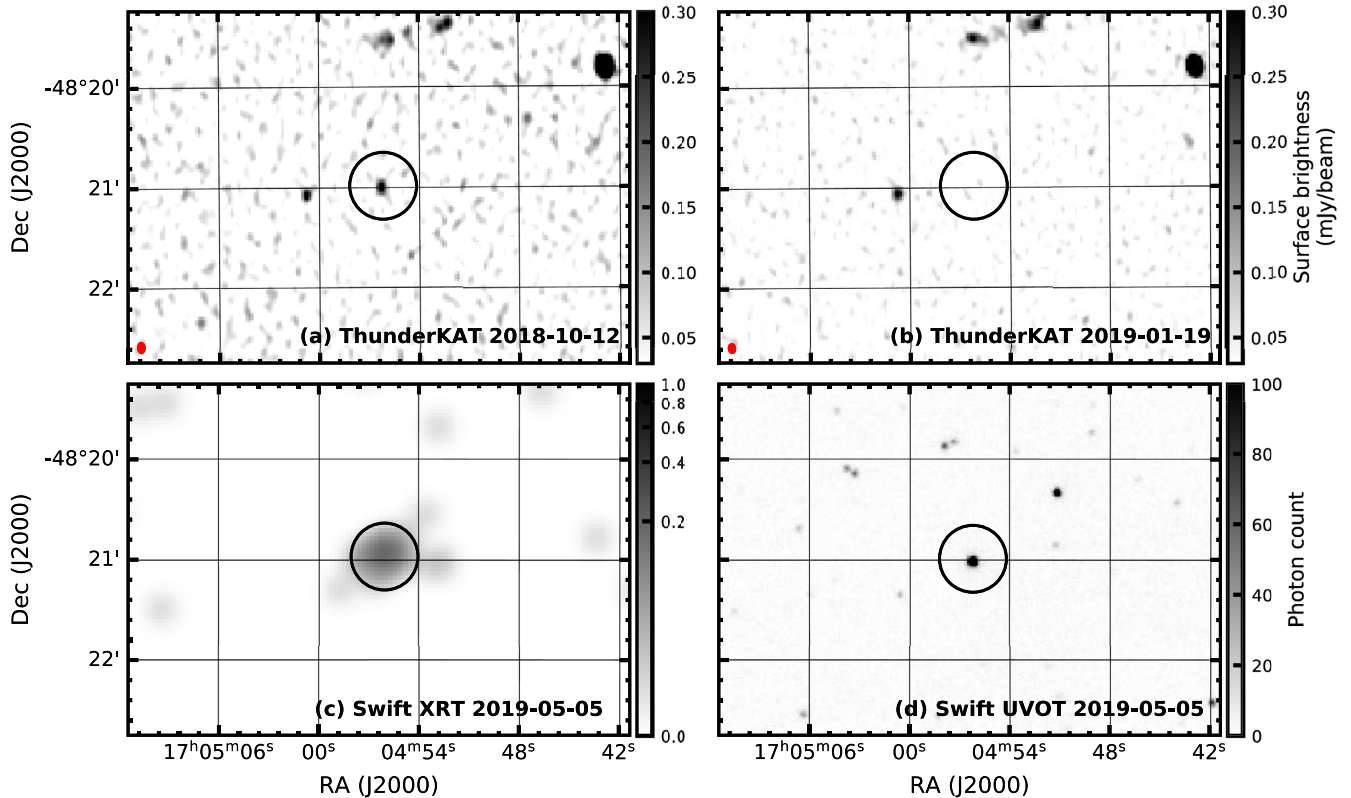


Figure 3. Radio, X-ray, and UV images with MKT J170456.2–482100 circled. Panels (a) and (b) show radio images from ThunderKAT. Panel (a) shows the source when it is detected at the highest flux density on 2018 October 12. Panel (b) shows an epoch, 2019 January 19, when the source is not detected. The synthesized beam is shown as a red ellipse in the bottom left corner of panels (a) and (b). Panels (c) and (d) respectively show the Swift XRT and Swift UVOT uvw1 images of the source from 2019 May 05. The grey scale on panels (a) and (b) shows the surface brightness in mJy beam^{-1} , the grey scale in panels (c) and (d) shows photon count.

Table 1. Summary of the archival optical photometry observations of TYC 8332-2529-1.

Survey	Optical band	Date of first observation	Date of last observation	Number of observations	Number of semesters
ASAS	V	2001 Jan 01	2009 Oct 12	534	9
KELT	V	2013 May 15	2015 Oct 13	3037	3
ASAS-SN	V	2016 Mar 10	2018 Sep 22	762	5
ASAS-SN	g	2018 Feb 26	Ongoing	645	3

3 IDENTIFICATION OF THE OPTICAL COUNTERPART TO MKT J170456.2–482100

The Python Blob Detector and Source Finder (PYBDSF; Mohan & Rafferty 2015) was used for improved astrometry to determine the position of MKT J170456.2–482100 in the epochs that it is detected. The position of the source was taken to be the mean of the J2000 positions: $17^{\text{h}}04^{\text{m}}56^{\text{s}}.2 - 48^{\text{d}}21^{\text{m}}00.45^{\text{s}}$ ($256.23450^{\circ} - 48.35012^{\circ}$) with standard deviations of 1.2 and 0.68 arcsec for the Right Ascension and Declination, respectively. We then searched for sources nearby to this position and found the star TYC 8332-2529-1, which has a *Gaia* data release 2 (Gaia Collaboration 2016; Bailer-Jones et al. 2018; Gaia Collaboration et al. 2018) ICRS J2015.5 position of $17^{\text{h}}04^{\text{m}}56^{\text{s}}.25 - 48^{\text{d}}21^{\text{m}}00^{\text{s}}.67$ with uncertainties of 0.044 and 0.037 arcsec, respectively, a proper motion of -3.694×10^{-3} and -6.963×10^{-3} arcsec yr^{-1} in Right Ascension and Declination, respectively, and a distance of 555 pc (68 percent confidence interval 541–568 pc). Taking proper motion into account, this position is 0.37 arcsec from the position of MKT J170456.2–482100, which means that we identify TYC 8332-2529-1 as the optical counterpart to MKT J170456.2–482100.

TYC 8332-2529-1 has been observed by numerous sky-surveys at various wavelengths. These include the original *Hipparcos* Tycho-2 catalogue (Høg et al. 2000), the Gunn *I* photometry from the Deep Near-Infrared Survey of the southern sky (DENIS; Fouqué et al. 2000), the Two-Micron All-Sky Survey (2MASS; Cutri et al. 2003), and the final *Wide-Field Infrared Survey Explorer* catalogue (*AllWISE*; Cutri et al. 2013). TYC 8332-2529-1 was classified as a main-sequence K-dwarf (K7V) by Pickles & Depagne (2010), and it was classified by the All-sky Automated Survey (ASAS; Pojmanski 1997) Catalogue of Variable Stars (ACVS; Pojmanski 2002) as a first-overtone delta Cepheid (DCEP_{F0}) or eclipsing binary with a period of 0.955 d. Richards et al. (2012) found, with a probability of ~ 50 percent, that the star is a type A Small Amplitude Red Giant (SARG; Wray, Eyer & Paczyński 2004) with a period of 21.245 95 d.

4 OPTICAL PHOTOMETRY

TYC 8332-2529-1 has been observed $\gtrsim 4800$ times in total over the last ~ 18 yr by ASAS (Pojmanski 1997), the Kilodegree Extremely Little Telescope (KELT; Pepper et al. 2007), and the All-Sky Automated Survey for Supernovae (ASAS-SN; Shappee et al. 2014; Kochanek et al. 2017). The details of the photometric observations of TYC 8332-2529-1 made by these surveys are shown in Table 1.

Once TYC 8332-2529-1 was identified as the source of MKT J170456.2–482100, we observed the star with the HIPPO Photopolarimeter (Potter et al. 2010) on the 1.9-m telescope at the South African Astronomical Observatory (SAAO), Sutherland, South Africa. HIPPO uses photomultiplier detectors, so is a photon counting instrument, obtaining measurements simultaneously in

two channels. Dual counter-rotating waveplates (1/2 and 1/4) modulate the signal at 10 Hz and allow simultaneous determination of all four Stokes parameters. The intrinsic time resolution is 1 ms for total intensity measurements (I) and 100 ms for the three other Stokes parameters (*Q*, *U*, and *V*). In practice data are accumulated over many 100 ms cycles, typically several 100 s, until a sufficient S/N is obtained. A log of these photopolarimetric observations is presented in Table 2.

4.1 Photometric variability

The automated ASAS pipeline classified TYC 8332-2529-1 as a 0.955 043 d variable. However, we performed a Lomb–Scargle (Lomb 1976; Scargle 1982; Perera et al. 2018) analysis of the data and found that the 0.955 043 d periodicity is due to the underlying periodicity of the observing cadence, instead of the intrinsic periodicity of the star. We found the highest precision period of the star to be 21.25 ± 0.04 d by performing a Lomb–Scargle analysis on the combined *V*-band data for the full ~ 18 yr of observations. This is consistent with the period of 21.246 d found by Richards et al. (2012) using machine-learning techniques on the ASAS observations.

We further investigated the period of TYC 8332-2529-1 by performing a Lomb–Scargle analysis on each semester⁶ of observations from ASAS, KELT, and ASAS-SN. We found that the period for each semester is consistent with 21.25 ± 0.04 d, as shown in Fig. 4. While the period over ~ 18 yr remains stable, Fig. 5 shows that the shape of the folded light curve and the phase of the peak varies over time. In some semesters the folded light curves show little or no variability and the period is less distinct. In the first semester of KELT observations (start MJD = 56427, 2013-03-15) the Lomb–Scargle analysis results in more power at half of the 21.25 ± 0.04 d period. Fig. 5 shows that this is due to a second peak in the folded light curve of this semester.

4.2 Spectral energy distribution

To obtain fundamental stellar parameters, the optical and infrared spectral energy distribution (SED) of TYC 8332-2529-1, shown in Fig. 6, was modelled. The SED was constructed from optical *B_T*- and *V_T*-band fluxes, obtained from the original *Hipparcos* Tycho-2 catalogue (Høg et al. 2000); the mean ASAS *V*-band magnitude (Pojmanski 1997); Gunn *I* photometry from DENIS (Fouqué et al. 2000); *J*, *H*, and *K_s* magnitudes from 2MASS (Cutri et al. 2003); and the [3.4] and [4.6] mag from *AllWISE* (Cutri et al. 2013). The [11.3] and [22] mag from *AllWISE* were not used, as TYC 8332-2529-1 lies 59 arcsec from IRAS 17011–4817, an optically obscured star that reaches second magnitude in these bands and provides considerable contaminating flux.

⁶We define a semester as the period in each year when the source is a night time source.

Table 2. Optical observation log for TYC 8332-2529-1, including the radial velocity measurements.

Telescope	Instrument	Wavelength range/filter (Å)	Resolving power	Date	UT start (hh:mm)	Exposure times	Radial velocity (v_{LSR} , km s $^{-1}$)
SAAO 1.9-m	HIPPO	3400–9000; no filter		2019-03-12	02:19	467 s	
					02:46	2020 s	
					02:48	2073 s	
					02:50	2065 s	
					02:51	1987 s	
SALT	HRS	3800–8800	15000	2019-02-25	02:42	1200 s	53 ± 3
				2019-03-24	00:16	1500 s	0 ± 3
				2019-04-17	22:45	1500 s	−30 ± 3
				2019-04-18	22:38	1500 s	−32 ± 3
				2019-06-08	19:18	1500 s	38 ± 3
				2019-06-10	19:30	1500 s	52 ± 3
SALT	RSS	6137–6953	4350	2019-03-01	01:55	600 × 5 s	20 ± 18
Cerro Tololo LCO	NRES	3800–8600	53000	2019-06-15	19:04	1200 s	24 ± 7
Sutherland LCO	NRES	3800–8600	53000	2019-06-01	01:20	1200 s	−35 ± 7
				2019-06-04	03:05	1200 s	−21 ± 7
				2019-06-06	02:08	1200 s	3 ± 7
				2019-06-09	03:34	1200 s	37 ± 7
				2019-06-14	03:06	1200 s	43 ± 7
				2019-06-19	07:20	1200 s	−16 ± 7
				2019-07-04	03:33	1200 s	49 ± 7
				2019-07-05	23:08	1200 s	43 ± 7

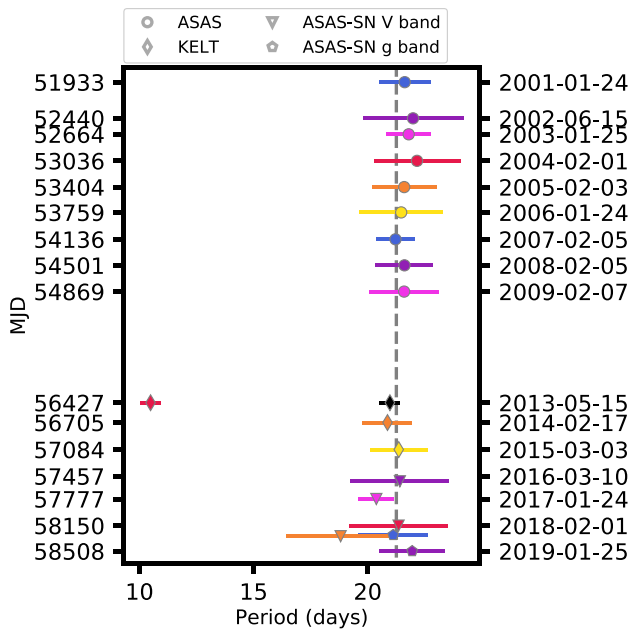


Figure 4. Periods found using the Lomb–Scargle algorithm for the different data sets divided into semesters, the MJD and date shown on the y-axis is the date of the first observation in the semester. The semesters are shown chronologically from top to bottom. The grey dashed line is the period found by performing a Lomb–Scargle on all of the V-band observations together: $P = 21.25 \pm 0.04$ d. In the first semester of KELT observations (start date: 2013 May 15) there was more power in the second harmonic of the period.

The software described in McDonald, Zijlstra & Boyer (2012) and McDonald, Zijlstra & Watson (2017) was used to fit BT-SETTL model atmospheres to the star. The scattering law of Draine (2003), assuming $R_V = E(B - V)/A_V = 3.1$, was used to correct the observed photometry for the effects of interstellar reddening,

with $E(B - V)$ as a free parameter. The *Gaia* Data Release 2 distance of 555 pc from Bailer-Jones et al. (2018) was adopted to scale the luminosity. A mass of $2.5 M_{\odot}$ was assumed to set the model surface gravity [$\log(g)$], but this has relatively little impact on the results.

Assuming a single stellar light source, a best fit was found for a temperature of $T_{\text{eff}} = 5519^{+405}_{-222}$ K and $E(B - V) = 0.51^{+0.08}_{-0.10}$ mag, giving a luminosity scaling of $L = 49.6^{+9.5}_{-8.7} L_{\odot}$. Unaccounted-for uncertainties include any intrinsic variability, contamination from other sources such as unresolved companions, or other departures from a simple model atmosphere, and a degree of correlation exists among the three parameters. The reduced χ^2 per degree of freedom is 0.93, hence the expected impact of these uncertainties is small.

To determine physical parameters from these values, stellar isochrones were generated using the Padova stellar evolution code⁷ (Marigo et al. 2017). The default settings were used, including the assumption of solar metallicity ($Z_{\odot} = 0.0152$). The star was identified as most likely being a sub-giant, completing its transition across the Hertzsprung gap to the base of the red giant branch. This period is characterized by a rapid inflation of the star towards the Hayashi limit as it becomes convective, following exhaustion of core hydrogen. Without a detailed metallicity derivation, a precise age and mass is impossible to ascertain. However, the age of the star is consistent with solar-metallicity isochrones in the age range $650 \text{ Myr} \lesssim t \lesssim 870 \text{ Myr}$, corresponding to a stellar mass of $2.22 M_{\odot} \lesssim M \lesssim 2.48 M_{\odot}$ and $\log(g) \approx 3.03 \pm 0.15$ dex. These values assume that the flux from any possible companion is negligible over the optical and near-IR SED, and that the star has taken an evolutionary path through the H–R diagram comparable to a single star with similar initial parameters.

⁷<http://stev.oapd.inaf.it/cgi-bin/cmd.3.2>

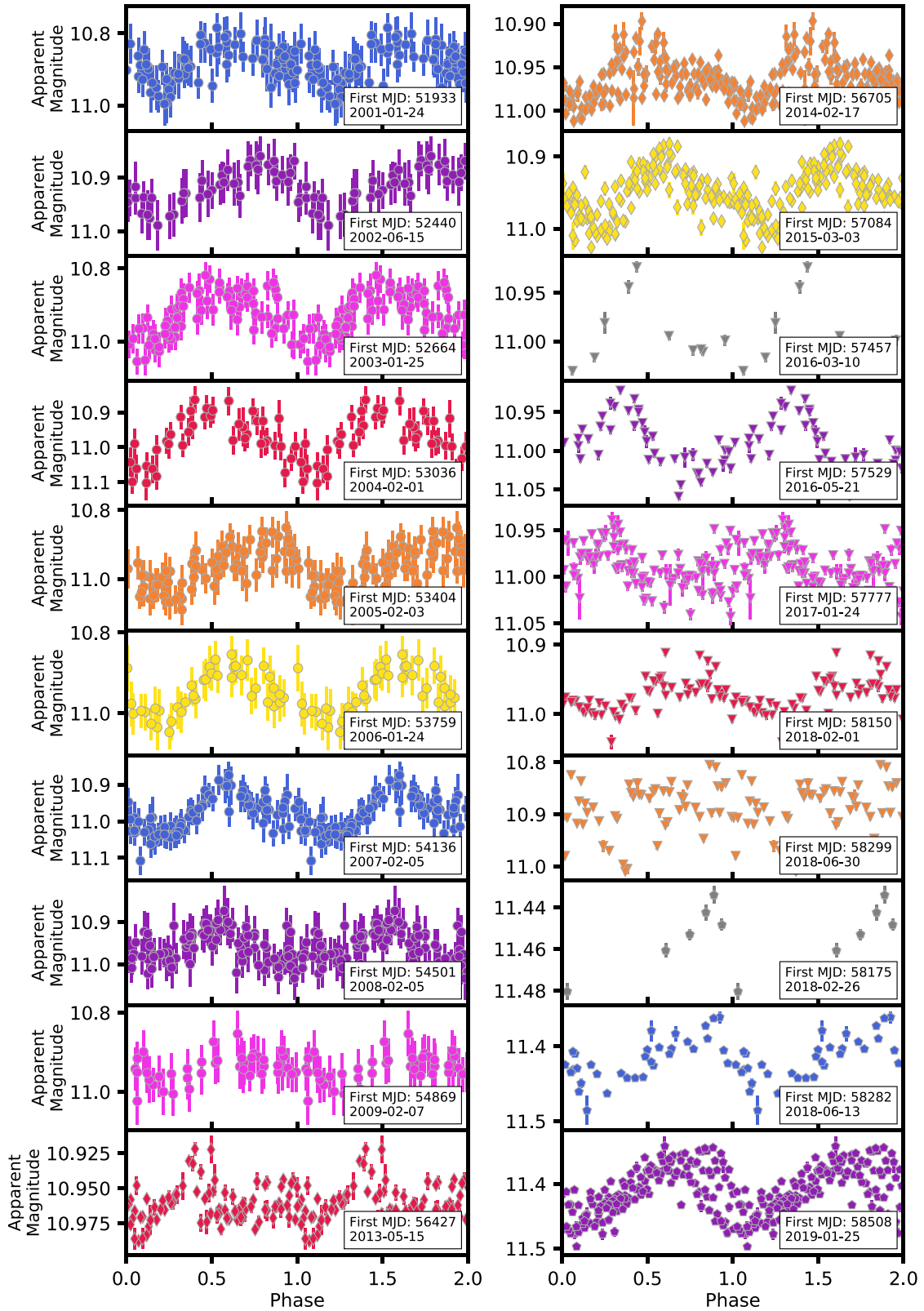
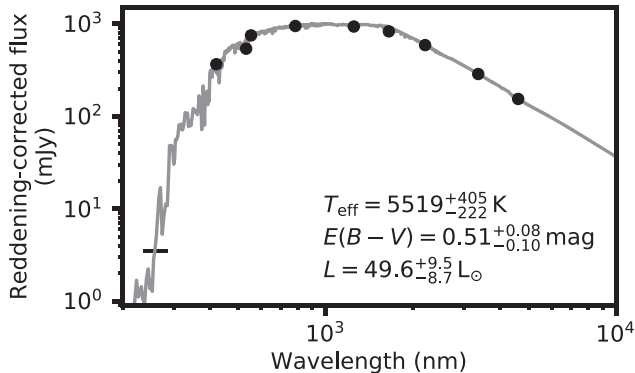
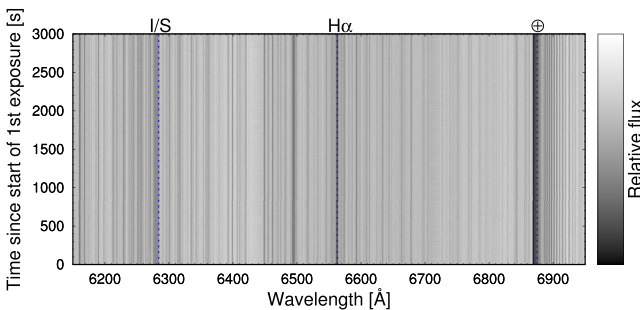


Figure 5. Optical light curves folded to a period of 21.25 ± 0.04 d. The observations are divided into semesters and have been folded and phase corrected to $\text{MJD} = 53571.1$. The MJD of the first observation of the semester is shown in the bottom right of each panel. The colours and symbols are the same as in Fig. 4, the grey points are the semesters that did not have enough points to perform a Lomb–Scargle analysis. The observations have been re-binned such that there is one measurement per day and outliers have been removed.

Table 3. Polarization measurements for TYC 8332-2529-1.

Filter	Linear (per cent)	θ ($^\circ$)	Circular (per cent)
<i>B</i>	2.56 ± 0.32	39.72 ± 4.44	-0.24 ± 0.22
<i>V</i>	2.87 ± 0.18	36.78 ± 2.93	0.00 ± 0.43
<i>R</i>	2.61 ± 0.08	39.21 ± 2.30	-0.04 ± 0.09
<i>I</i>	2.33 ± 0.16	41.11 ± 1.59	0.01 ± 0.12

**Figure 6.** SED of TYC 8332-2529-1, showing literature sources used in the fit (black data points, see text). The black line shows our Swift UVOT observation, with length of the line denoting the width of the filter. The grey spectrum shows the best-fitting BT-SETTL model, with the parameters shown.**Figure 7.** 600×5 s SALT RSS red spectra, covering 6137–6953 Å, of TYC 8332-2529-1 taken on 2019 March 1. The absorption lines include the narrow H α line (right of centre), the telluric oxygen *B* band (far right, marked with \oplus) and many other photospheric lines. Each spectrum was normalized to account for slit losses.

4.3 Photopolarimetry

TYC 8332-2529-1 was observed using HIPPO on 2019 March 12. The results show that TYC 8332-2529-1 was not varying significantly photometrically over the ~ 2 h observation period and furthermore show no evidence for higher frequency variability over timescales of ~ 1 –1000 s. An initial filterless observation was performed followed by five repeat *BVRI* filter sequences. There was no evidence of variations in the polarization between different filter sequences and the average polarization parameters for the *BVRI* filters are presented in Table 3. The level of linear polarization measured is consistent with the values expected for the ISM at the Galactic coordinates for TYC 8332-2529-1.

5 OPTICAL SPECTROSCOPY

The variable nature of the optical light curve of TYC 8332-2529-1 is suggestive of an active flare star or binary companion. We therefore undertook a spectroscopic campaign to further understand the nature of the variability. We observed TYC 8332-2529-1 with the Southern African Large Telescope (SALT; Buckley, Swart & Meiring 2006) situated at the SAAO using the High Resolution Spectrograph (HRS; Bramall et al. 2012; Crause et al. 2014) and the Robert Stobie Spectrograph (RSS; Burgh et al. 2003). HRS is a fibre-fed, dual-beam, white pupil, vacuum-stabilized high resolution ($R = 15\,000$ –80 000, depending on mode and wavelength) spectrograph, while RSS is a prime focus low to medium resolution ($R = 350$ –6500, depending on grating/angle/slit choice) slit spectrograph. TYC 8332-2529-1 was also observed by the Las Cumbres Observatory (LCO; Brown et al. 2013) Network of Robotic Spectrographs (NRES). The NRES network uses high-resolution ($R \sim 53\,000$) optical echelle spectrographs on two, 1 m telescopes simultaneously with a ThAr calibration source. The details of the spectroscopic observations are summarized in Table 2.

5.1 High-resolution spectroscopy

Spectra of TYC 8332-2529-1 were taken with the SALT HRS in the low-resolution (LR) mode on six occasions, shown in Table 2, in clear conditions with ~ 1 arcsec seeing. Both blue (3800–5550 Å) and red (5450–9000 Å) spectra were obtained and reduced using the weekly set of HRS calibrations, including ThAr arc spectra and QTH lamp flat-fields.

The primary reduction was conducted using the SALT science pipeline, PYSALT (Crawford et al. 2016), correcting for overscan, bias and gain correction. The spectral reductions were then undertaken using a MIDAS-based echelle reductions package (see details in Kniazev, Gvaramadze & Berdnikov 2016). For the blue spectra, the average S/N had a minimum of ~ 10 for the bluest orders (~ 3980 –4000 Å), which increased linearly to ~ 120 at 4600 Å, and reached a maximum of ~ 140 at 5400 Å. For the red spectra, the S/N increased linearly from ~ 80 at 5550 Å to ~ 180 at 7000 Å, peaking at ~ 200 by 8000 Å and then declining slightly to ~ 160 by 8800 Å. This behaviour is consistent with the late spectral type (K4-5) attributed to TYC 8332-2529-1.

5.2 Fast time-resolved spectroscopy

High time resolution spectroscopy of TYC 8332-2529-1 was obtained on 2019 March 1 using the RSS in frame transfer mode in clear conditions with 1 arcsec seeing. The PG2300 grating was used at a grating angle of 49.25° , which gave a wavelength coverage of 6137–6953 Å at a mean resolution, with a 1.5 arcsec slit, of 1.5 Å. A set of 600 continuous 5 s exposures were obtained, with no dead time between frames because of the frame transfer mode. The data were reduced using the PYSALT package (Crawford et al. 2016) to make the usual corrections for overscan, bias, gain, and crosstalk. A 2 s neon arc lamp exposure, taken after the observations, was used to wavelength calibrate the spectra. The IRAF package was then used to perform the spectral reductions, using the tasks in `twodspec`. In Fig. 7 we show a 2D spectral image comprising of all 600 spectra, extending from 6137 to 6953 Å and stacked sequentially in time, increasing upwards.

Little variation is seen in the spectra, in terms of line strength or position, and certainly no evidence of H α variability.

5.3 Interpretation

5.3.1 Radial velocity curve

The wavelength calibration of the first SALT/HRS spectrum was initially checked against a telluric-feature spectrum in the regions around 6000 and 6800 Å, and a good match was found. The radial velocity amplitude of the target star was then derived by cross-correlating the spectrum with a BT-SETTL model atmosphere ($T_{\text{eff}} = 5000$ K, $\log g = 4.0$ dex, solar metallicity; Allard et al. 2003) in the region 6100–6240 Å, chosen to be largely free of telluric absorption or emission features. Radial velocities for subsequent spectra were checked in a similar manner, then cross-correlated against this first spectrum, to ensure a more accurate relative velocity. Radial velocities for the LCO spectra were corrected from the pipeline-reduced version by -88 km s $^{-1}$. Observed radial velocities were corrected to the Local Standard of Rest using the STARLINK tool RV (Currie et al. 2014). The resulting radial velocities are listed in Table 2 and shown in Fig. 8. Formal errors are dominated by the accuracy of the wavelength calibration, and approximate uncertainties have been assigned on this basis. Since these cross-instrument velocity uncertainties dominate the radial velocity amplitude, and since the photometric period is well known, we do not perform a formal fit to the radial velocity data.

5.3.2 Spectral characterization

The spectrum of TYC 8332-2529-1 appears consistent with the K giant determined from the SED modelling in Section 4.2. The Ca II H and K lines are in emission reaching the continuum level and there is infilling of the infrared triplet, which indicates that the star is chromospherically active (the Ca λ 8662 line is shown in Fig. 9).

There are no emission lines in the optical spectrum other than the chromospheric emission of the H and K lines and the infilling of the infrared triplet. This indicates a general lack of hot material in the vicinity of the unseen companion, as emission lines are generally seen even in relatively faint accretion discs (e.g. RW Sextantis; Hernandez et al. 2017).

5.3.3 Spectral variation

Despite showing considerable radial velocity variations, the shape of the red portion of the TYC 8332-2529-1's spectrum remains mostly constant, generally to within ~ 1 per cent (Fig. 9). Most variations appear to be caused by telluric features, interstellar gas lines and diffuse interstellar bands (DIBs). Subtle, percentage-level variations in the immediate vicinity of strong atomic lines can be traced to variation in line widths, likely due to very small variations in the recorded spectrum caused by either seeing or telescope focus.

However, variations in the line wings of bright lines exist. In the difference spectrum, this results in negative and positive features offset from the stellar velocity. The systemic rest velocity ($\sim \pm 43$ km s $^{-1}$) is marked with purple arrows in Fig. 9. In stronger and bluer lines, extra absorption is seen beyond the systemic rest velocity, indicating the absorbing material is on the opposite side of the system's barycentre and hotter than the giant star. Modelling these components as one would a companion star gives a reflex velocity amplitude of $K_2 \approx 12 \pm 10$ km s $^{-1}$. The relative line depths suggest a spectral temperature of order 7000 K, but this is very approximate as the K-giant star's spectrum cannot be accurately subtracted with our current data.

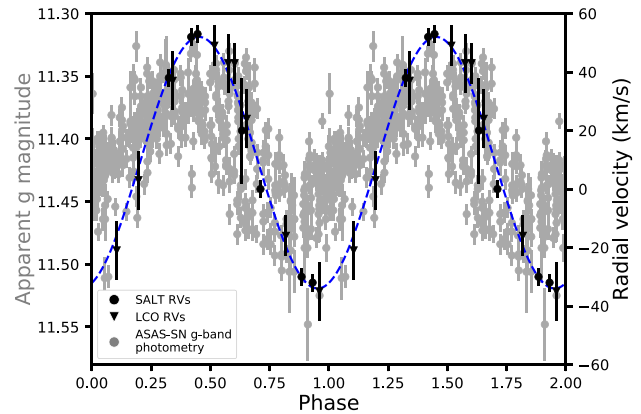


Figure 8. Radial velocity (black) and optical photometry (grey) curves, phased to a period of 21.25 ± 0.04 d and an epoch of $\text{MJD} = 53571.1$. The blue dashed line represents a visual, sinusoidal fit to the radial velocity points.

6 RADIO PULSATION SEARCHES

As MKT J170456.2–482100 is coincident with a star in a spectroscopic binary, it is possible that the companion star may be a pulsar or pulsating white dwarf, similar to AR Scorpii (Marsh et al. 2016). To investigate this possibility, we performed high-time resolution observations to search for pulsations over a range of periods spanning milliseconds to minutes. We observed the position of MKT J170456.2–482100 on UTC 2019 January 30 with the ultra-wide-bandwidth (UWL) receiver deployed on the 64-m Parkes radio telescope in Australia. The output of the receiver is processed by the Medusa Graphics Processing Unit (GPU) cluster which produces 8-bit, full-Stokes filter banks with spectral and temporal resolutions of 1 MHz and 128 μ s respectively over the 3328 MHz band from 0.740 to 4.03 GHz. The data were recorded in PSRFITS format and processed using PSRCHIVE tools (Hotan, van Straten & Manchester 2004). Since our data were strongly affected by radio frequency interference (RFI) we used the `clfd`⁸ package described in (Morello et al. 2019) to perform more sophisticated RFI mitigation. Since radio waves are dispersed by free electrons in the interstellar medium along the line of sight, the data need to be corrected for this dispersion delay before any analyses can be made. This dispersion delay can be quantified by the dispersion measure (DM), which is the integrated free electron column density along the line of sight. We dedispersed the data over a range of trial DMs, $0.0 \leq \text{DM} \leq 30.0$ pc cm $^{-3}$ in steps of 1 pc cm $^{-3}$ to account for the uncertainty in the distance to TYC 8332-2529-1. The DM range was estimated using various distances to the source (Bailer-Jones et al. 2018), and the NE2001 model (Cordes & Lazio 2002). For each trial DM, the resulting dedispersed time series was searched for long and short-period pulsations using both a Fast Folding Algorithm (FFA) and a Fast Fourier Transform (FFT), respectively.

An FFT was performed using the PRESTO suite of pulsar search and analysis software (Ransom, Eikenberry & Middleton 2002). For each trial DM time series resulting from the dedispersion mentioned above, we performed an FFT using the REALFFT routine. We then used the ACCELSEARCH routine which sums 16 harmonics incoherently for each frequency bin to improve the signal-to-noise (S/N), and every detection above a S/N of 5 was saved for further

⁸<https://github.com/v-morello/clfd>

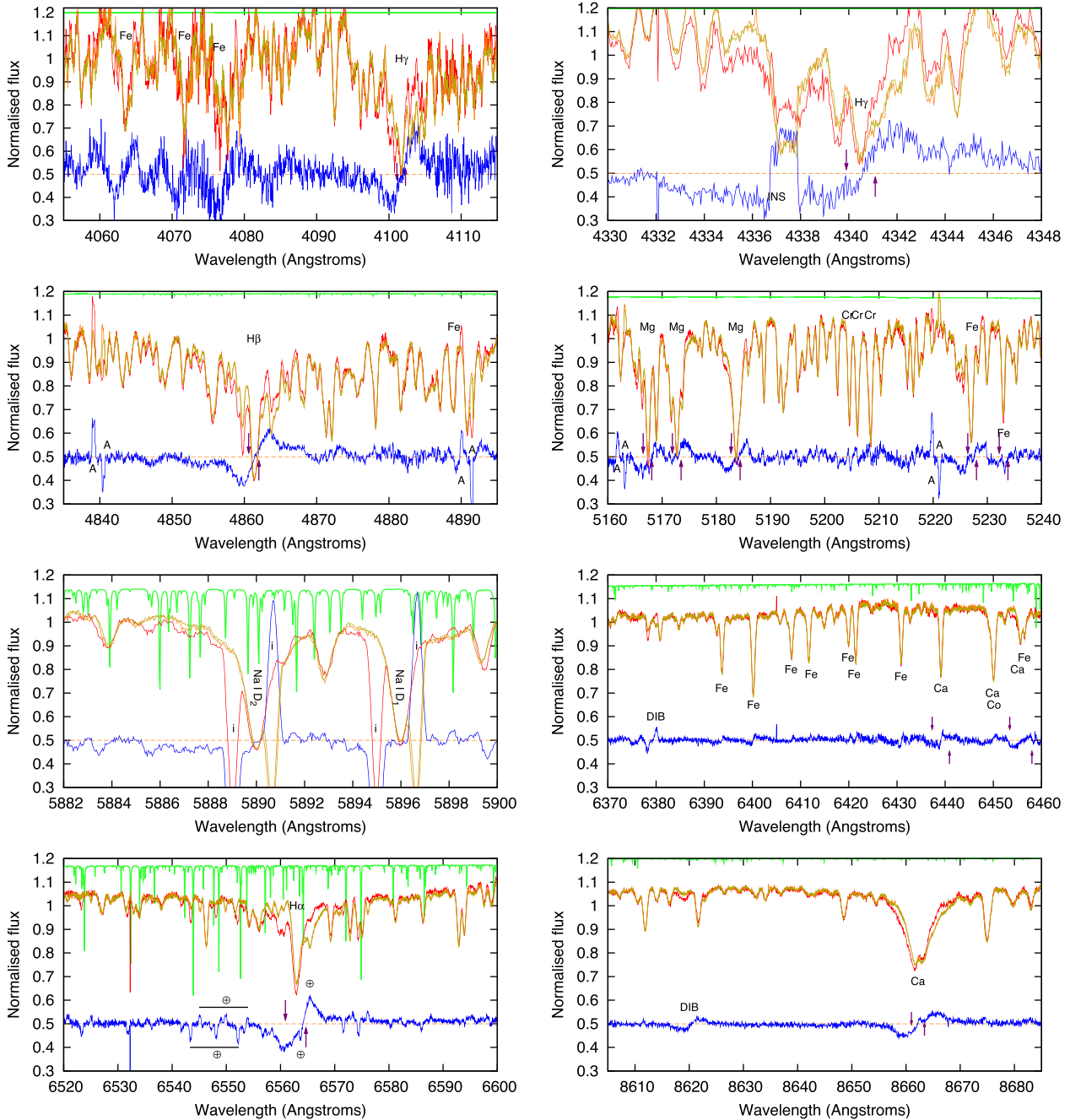


Figure 9. SALT/HRS spectra, flux-normalized and corrected to the radial velocity amplitude of the giant star. The first, second, and fourth SALT/HRS spectra from Table 2 are shown in red, orange, and gold. Of the spectra taken, these represent those with the most-positive, closest-to-average, and most-negative radial velocities. The light green spectrum (top) shows the telluric absorption at rest velocity, offset vertically by +0.2. The difference between the first and fourth spectra is shown in blue (bottom), offset by +0.5 in normalized flux. Instrumental artefacts (INS), telluric absorption (\oplus) and airglow emission (A), and interstellar (i, DIB) lines are visible in this differenced spectrum. Purple arrows represent the systemic rest velocity and are explained in the text in Section 5.3.3. The minima/maxima of the blue difference spectrum are seen $\pm \sim 12 \text{ km s}^{-1}$ beyond these arrows.

inspection. We folded the time series for each candidate period using the PREPFOLD routine and the folded time series was visually inspected to check whether it resembled a true astrophysical source. No significant periodic pulsations were detected above the threshold S/N.

While FFTs are a standard technique used for searching pulsars and short-period sources (e.g. Rajwade et al. 2018), they become less sensitive when searching for longer periods as red noise

becomes a dominant factor. Therefore, we used the RIPTIDE⁹ FFA algorithm developed by Morello et al. (in prep) to search for periods ranging from 1 s to 10 min. Similar to the FFT, we dedispersed and folded each time series and vetted the candidates for significant pulse profiles above an S/N of 8. We did not detect any significant periodic pulsations above the S/N threshold.

⁹<https://bitbucket.org/vmorello/riptide/src/master/>

Table 4. Details of the UV observations of MKT J170456.2–482100 taken with the Swift UVOT instrument with the UVW1 filter (centre wavelength 2600 Å).

Date	Time (UTC)	Exposure (s)	AB magnitude
2019-04-18	20:45:17	393.93	15.02 ± 0.02
	22:18:47	1012.48	15.01 ± 0.02
2019-05-05	03:30:48	500.25	15.00 ± 0.02
	05:11:06	221.68	14.98 ± 0.03
	09:41:44	306.54	15.00 ± 0.02

In addition to the periodicity searches described above, we also performed offline single pulse searches on the data using the GPU based HEIMDALL¹⁰ single pulse search software. The HEIMDALL code reads in a subset of the data at a time, and processes it over the DM-width parameter space in search of single pulses. We expect any single pulse to be detected over a range of DMs close to the true value. A clustering algorithm merges these events into one before returning the DM that yields the maximum S/N. This reduces the number of candidates considerably as thousands of events can be identified prior to merging, which will then eventually result in a much lesser number of unique candidates. The data were searched for bursts in DM and width space by dedispersing over trial DMs in the range $0.0 \leq \text{DM} \leq 30.0 \text{ pc cm}^{-3}$ and convolving with a series of sliding boxcar filters (square pulse) of width $2^0 \leq W \leq 2^{17}$ time samples, respectively. No significant bursts were detected above an S/N threshold of 10.

7 UV OBSERVATIONS

We were granted two target of opportunity observations of MKT J170456.2–482100 on 2019 April 18 and 2019 May 05 with the Swift UVOT (Mason et al. 2000) and XRT (Hill et al. 2005) instruments. For the Swift UVOT observations we used the UVW1 filter (centre wavelength 2600 Å). Two exposures of 393.9 and 1012.5 s were taken on 2019 April 18 and three exposures of 500.25, 221.68, and 306.54 s were taken on 2019 May 05. The magnitudes of the source are shown in Table 4, and the UV is plotted on the SED of TYC 8332-2529-1 in Fig. 6. As can be seen in Table 4, there is no evidence of variability greater than 0.02 mag in the UV from these observations on time-scales of a few hours and 27 d.

8 X-RAY OBSERVATIONS

The simultaneous Swift XRT (Hill et al. 2005) observations on the 2 d were 1.4 and 1.0 ks, respectively. We used the XRT photon counting mode in the 0.3–12 keV band. During the first observing session we detected 18 events from the source. The probability that these counts are from the background is 1.1×10^{-61} . In the second session, 19 events were detected from the source with a probability of 1.54×10^{-57} that the photons are from the background. Based on these observations we claim a faint X-ray detection of MKT J170456.2–482100. Most events detected from the source were in the soft part of the band, 0.3–3 keV. Due to the low photon count we were unable to fit a spectrum to the individual observations, but we were able to fit a power law to the combined spectrum from both epochs to obtain a flux density of $7.3_{-2}^{+3} \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$. The parameters for this fit are shown in Table 5 and one of the X-ray images is shown in Fig. 3.

Table 5. Parameters of a power-law fit to the XRT spectrum of MKT J170456.2–482100 with 1σ error bars in parentheses. Γ is the power-law photon index, N_{H} is the neutral hydrogen column density along the line of sight. F_{Xabs} is the absorbed X-ray flux and χ^2 is the reduced chi-squared statistic of the fit.

Parameter	Value
Γ	2.4 (+0.8, −0.7)
$N_{\text{H}} \text{ (cm}^{-2}\text{)}$	$4.0 \text{ (+4.0, −3.0)} \times 10^{21}$
$F_{\text{Xabs}} \text{ (ergs cm}^{-2} \text{ s}^{-1}\text{)}$	$7.3 \text{ (+3, −2)} \times 10^{-13}$
χ^2	2.3

9 DISCUSSION

MKT J170456.2–482100 is a transient radio source coincident with a spectroscopic binary where one of the objects is the chromospherically active K-type giant TYC 8332-2529-1 (spectral type K4-5). In the radio we see a bright event occur in 2018 October, followed by decreasing radio brightness with some underlying variability. We do not find any radio pulsations. MKT J170456.2–482100 was detected in both the UV and the X-ray.

The approximately 18 yr of optical photometry shows that TYC 8332-2529-1 varies with a period of $21.25 \pm 0.04 \text{ d}$, but that the shape and amplitude of the light curve changes over time. The optical spectra show that MKT J170456.2–482100 is a spectroscopic binary with a period matching the photometric period. The line-of-sight radial velocity amplitude of TYC 8332-2529-1 is $\pm 43 \text{ km s}^{-1}$. There are no emission lines in the spectra, as would be expected for an accretion flow or accreting compact companion. The overall spectrum of the star is highly stable, but strong lines show small-scale variation consistent with an absorption component in orbital antiphase with the K-giant. Blue lines are preferentially affected, and the combination is consistent with a fainter, broader-lined object on the other side of the barycentre, with a line-of-sight radial velocity amplitude of $12 \pm 10 \text{ km s}^{-1}$ and a temperature of $\sim 7000 \text{ K}$. While the variations are clear, we are not able to extract this signal with sufficient clarity to state whether we have detected the signature of a spectroscopic binary companion, or some other physical effect within the system. Further spectral monitoring at high signal-to-noise will be necessary to obtain better phase coverage.

In Fig. 10 we plot the ASAS-SN semesters beginning on 2018 June 13 (MJD: 58282), 2018 June 30 (MJD: 58299), and 2019 January 25 (MJD: 58508) as well as the radio light curve from the MeerKAT observations. There is no clear periodicity in the 2018 June 30 semester of ASAS-SN V-band observations, which coincides with the bright radio event that occurred in 2018 October. However, in the 2019 January 25 semester of ASAS-SN g-band observations the periodicity has returned, with increasing amplitude over time. This is coincident with the decreasing radio flux over time. This could indicate that the cause of the optical photometric period was disrupted by the same event that caused the radio flaring in 2018 October. In Fig. 5 we can also see that the light curve is less defined in the 2013 May 15 (MJD: 56427) and 2014 February 17 (MJD: 56705) semester, but that the periodicity returns in the 2015 March 03 (MJD: 57084) semester. This could imply that a similar outburst to the 2018 October event occurred before or during the 2013 May 15 semester. A search of archival radio data has not revealed any data at these times to see whether there was a corresponding radio outburst. We also see, in Fig. 11, that the optical brightness of the star and the amplitude of the variability varies over time. Particularly that the 2018 June 30 semester of ASAS-SN V-band observations, that ended in 2018 October during the radio

¹⁰<https://sourceforge.net/projects/heimdall-astro/>

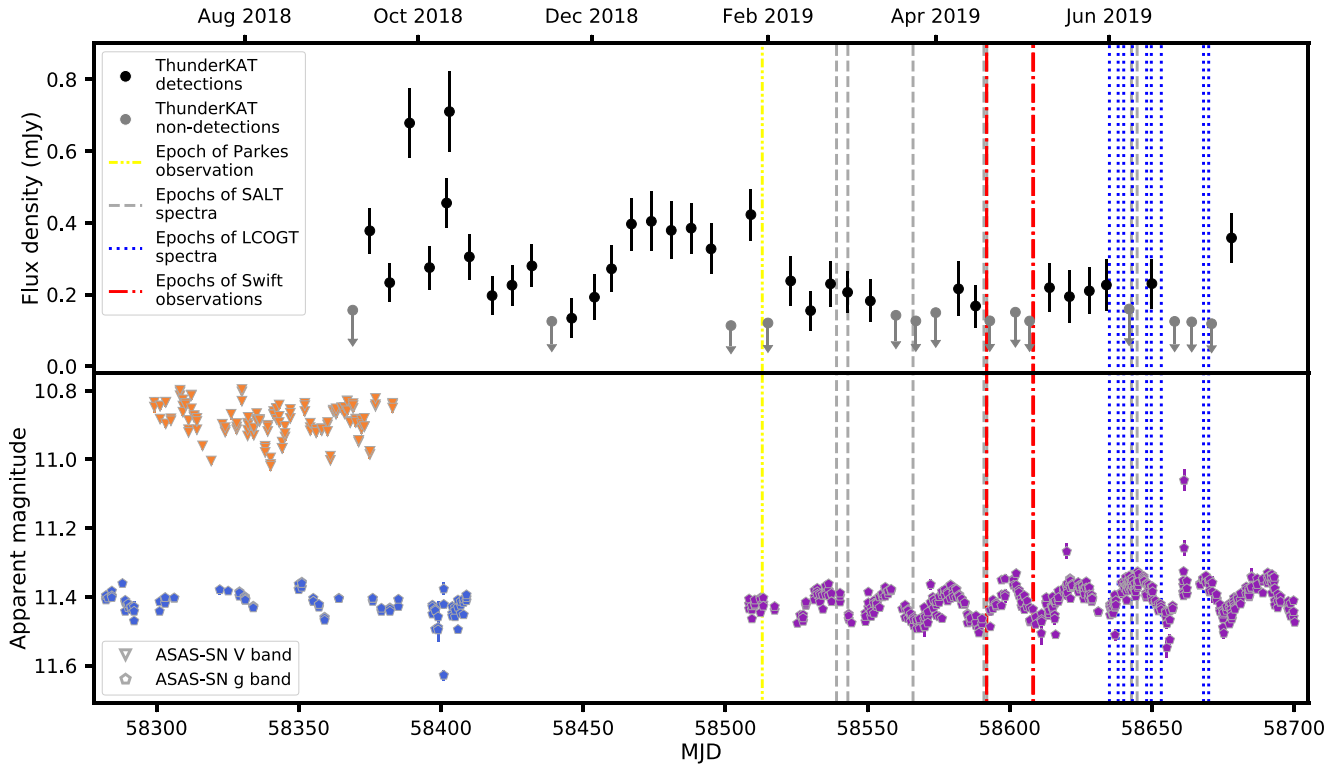


Figure 10. MeerKAT radio flux densities (top panel) and the ASAS-SN semesters (bottom panel) beginning on 2018 June 13 (MJD: 58282, *g* band), 2018 June 30 (MJD: 58299, *V* band), and 2019 January 25 (MJD: 58508, *g* band), the colours and symbols are the same as in Figs 4 and 5. The epochs of the LCO and SALT spectra, the *Swift* observations and the Parkes observations are plotted as vertical lines in both panels.

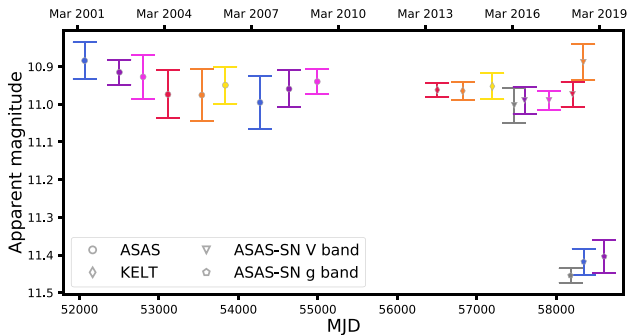


Figure 11. The weighted mean of each semester of optical observations, the bars indicate the standard deviation of the semester. The same colours and symbols are used here as in Figs 4 and 5.

flare, are significantly brighter than previous semesters of optical photometry. We also note that there appears to be a flare in the ASAS-SN *g*-band observations in 2019 June; however, this is an instrumental issue as the stars nearby to TYC 8332-2529-1 show a similar ‘flare’ in the same epoch in their light curves.

The variable shape of the photometric light curve could indicate the presence of large star spots on the surface of the K star (e.g. Frasca et al. 2005; Biazzo et al. 2006). For example, the shape of the *V*-band light curve in the latest semester of ASAS-SN *g*-band observations could be caused by the presence of two large star spots, similar to IM Peg (Biazzo et al. 2006). The variability in the shape and phase of the light curve over time could indicate that the spots change in size, number, and position on the star’s surface. Assuming

Table 6. The covering fraction, $f_{S,\min}$, of spots on the K-star assuming different values for c , the spot contrast. These values were calculated using the ASAS-SN *V*- and *g*-band observations. The semester start date is the date of the first observation of that semester.

Band	Semester start date	$f_{S,\min}$			
		$c = 0.0$	$c = 0.3$	$c = 0.5$	$c = 0.9$
<i>V</i>	2016-03-10	0.12	0.17	0.24	1.21
<i>V</i>	2016-05-21	0.07	0.09	0.13	0.66
<i>V</i>	2017-01-24	0.18	0.27	0.37	1.87
<i>V</i>	2018-02-01	0.14	0.20	0.27	1.37
<i>g</i>	2018-02-26	0.05	0.08	0.11	0.54
<i>g</i>	2018-06-13	0.17	0.24	0.34	1.70
<i>V</i>	2018-06-30	0.11	0.16	0.23	1.14
<i>g</i>	2019-01-25	0.15	0.22	0.31	1.55

that the variability is caused by star spots, we can calculate the covering fraction, i.e. the fraction of the star covered by star spots using (Morris et al. 2018):

$$f_{S,\min} = \frac{1 - \min \mathcal{F}}{1 - c} \quad (1)$$

where \mathcal{F} is the flux of the star divided by the mean flux and c is the spot contrast. A spot with the same intensity as the photosphere of the star would have $c = 0$, while a completely dark spot would have $c = 1$ (Morris et al. 2018). Sunspots have a contrast value of $c = 0.3$. The covering fractions for the ASAS-SN *V*- and *g*-band semesters are shown in Table 6, suggesting that $\gtrsim 10$ per cent of the stellar disc is spotted as a temporal average. If TYC 8332-2529-1 is spotted,

the radio activity and X-ray observations could also be from the star, similar to an RS CVn system. Guedel & Benz (1993) showed that there is a correlation between the average soft X-ray luminosity and radio luminosity for active stars. If we place the average radio luminosity and X-ray luminosity of MKT J170456.2–482100 on the Benz–Güdel plot in Fig. 12, then it fits well with the known values for RS CVn systems.

The minimum mass for a star to have evolved into a sub-giant in 13.8 Gyr is $0.935 M_{\odot}$ (e.g. Marigo et al. 2017; Tang et al. 2014; Chen et al. 2015, 2014). TYC 8332-2529-1 is not significantly metal-poor, but to take into account possible metallicity effects (McDonald & Zijlstra 2015) we use a conservative minimum mass for the star of $0.8 M_{\odot}$. We know from the spectra that the velocity amplitude of TYC 8332-2529-1 is 43 km s^{-1} , and if we assume, from the sinusoidal shape of the radial velocity curve shown in Fig. 8, that the orbit is circular we can then use the mass function to calculate the mass of the companion (m_2):

$$\frac{m_2^3}{(m_1 + m_2)^2} \sin^3 i = \frac{P}{2\pi G} v_1^3, \quad (2)$$

where m_1 and v_1 are the mass and radial velocity amplitude of TYC 8332-2529-1, respectively, i is the inclination of the binary system, and P is the orbital period. Assuming an inclination of 90° , a period of $21.25 \pm 0.04 \text{ d}$ from the optical photometry, and a K-star mass of $0.8 M_{\odot}$ gives a minimum companion mass of $\sim 0.75 M_{\odot}$ and a semimajor axis of $\sim 0.17 \text{ au}$. If the mass of TYC 8332-2529-1 is between ~ 2.22 and $\sim 2.48 M_{\odot}$ (Section 4.2), then an inclination of 90° gives a minimum companion mass of $1.29 M_{\odot}$. These values show that the companion may be a white dwarf.

However, if we interpret the spectral variations we see as absorption lines from a companion star on the opposite side of the barycentre to TYC 8332-2529-1, its low radial velocity amplitude ($\sim 12 \pm 10 \text{ km s}^{-1}$; see Section 5.3.3) requires some explanation. The binary companion to TYC 8332-2529-1 does not significantly affect the shape of the SED (Section 4.2; Fig. 6). The blue and UV photometry in particular mean that it is intrinsically much fainter and/or has a temperature less than $\sim 7000 \text{ K}$. If the companion is a main-sequence star, this combination restricts its mass to $\lesssim 1.5 M_{\odot}$.

Still assuming that the $\sim 12 \pm 10 \text{ km s}^{-1}$ is from a companion to TYC 8332-2529-1, we can take the ratio of the velocities to find a mass ratio. We then find that the mass of the companion is between $\sim 2 \times M_{\text{K giant}}$ and $\sim 22 \times M_{\text{K giant}}$. Taking our conservative minimum mass of TYC 8332-2529-1, $\sim 0.8 M_{\odot}$, provides a companion mass between ~ 1.6 and $\sim 17.2 M_{\odot}$, which corresponds to a semimajor axis, for a circular orbit, of 0.20 and 0.40 AU, respectively. Our modelled mass of TYC 8332-2529-1 $2.22\text{--}2.48 M_{\odot}$, would mean a companion mass between ~ 4.3 and $\sim 53.3 M_{\odot}$, which corresponds to a semimajor axis of 0.28 and 0.57 au, respectively. These masses are not compatible with a white dwarf, or with a star with a temperature less than $\sim 7000 \text{ K}$, and would lead to significant tidal distortion of TYC 8332-2529-1, which would be visible in the optical light curve.

One possible explanation is that MKT J170456.2–482100 is a triple-system, where the companion to TYC 8332-2529-1 is actually two objects orbiting each other, with the semimajor axis (a_{inner}) of the inner orbit much smaller than the semimajor axis (a_{outer}) of the outer orbit. If we assume that the inner orbit is circular with an inclination of 90° , and is comprised of two objects of equal mass (and hence equal radial velocity), we can calculate the minimum radial velocity amplitude of those objects. If we assume the minimum masses from above ($M_{\text{K giant}} = 0.8 M_{\odot}$, $M_{\text{companion}} = 1.6 M_{\odot}$, $a_{\text{outer}} = 0.2 \text{ au}$), and a semimajor axis of < 10 per cent of

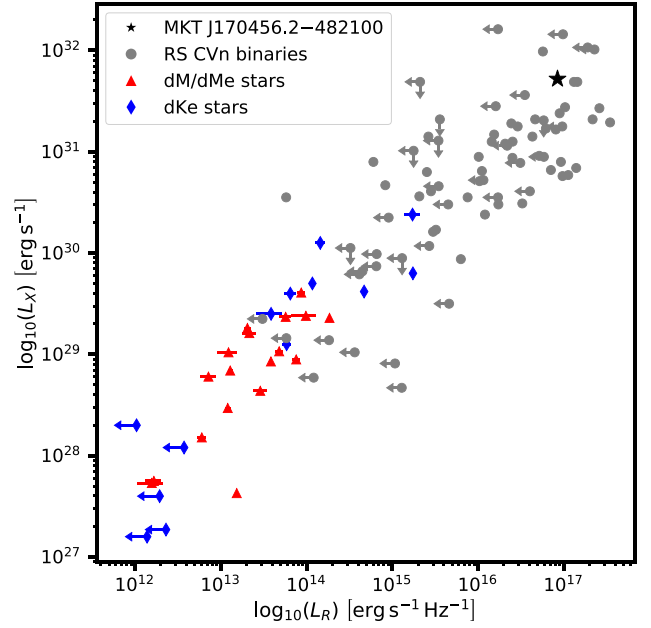


Figure 12. Plot of the average soft X-ray and radio luminosity of active stars from Guedel & Benz (1993), Benz & Guedel (1994), and references therein. dM, dMe, and dKe stars are magnetically active stars. The uncertainties were not available on many of these archival data points, but are included where possible. The X-ray luminosity for MKT J170456.2–482100 was found by fitting a spectrum to all of the Swift XRT observations we have obtained, for the Swift 0.3–12 keV band. The MKT J170456.2–482100 radio luminosity is the weighted mean of the radio luminosity across all epochs. There are uncertainties on the MKT J170456.2–482100 values, but they are too small to be seen.

the outer orbit ($a_{\text{inner}} = 0.02 \text{ au}$), then the velocity amplitude of one object in the inner orbit is $\gtrsim 130 \text{ km s}^{-1}$. This is inconsistent with a radial velocity amplitude of $12 \pm 10 \text{ km s}^{-1}$, and the spectral variations we see would reflect these higher velocity amplitudes.

Further investigation into this system is required to determine the nature of the companion to TYC 8332-2529-1, as well as to confirm the origin of the X-ray, radio, and UV flux. As TYC 8332-2529-1 is chromospherically active, the X-ray, radio, and UV emission may be from this star. A UV spectrum would confirm the nature of the UV emission. This MeerKAT field is being observed weekly by ThunderKAT until 2023 September, and ASAS-SN continues to observe this field. This will help to determine whether the radio and optical variability is related, which will help determine whether the radio activity is from flaring activity on TYC 8332-2529-1. We are also observing MKT J170456.2–482100 weekly in the X-ray for 9 weeks with *Swift*. This will help determine whether the source is variable in the X-ray. *TESS* observations of the source will be released later in 2019, which may reveal shorter time-scale optical variability.

10 CONCLUSIONS

We report the discovery of MKT J170456.2–482100, the first transient discovered with MeerKAT. MKT J170456.2–482100 coincides with the K4-5V type sub-giant star TYC 8332-2529-1. Using $\sim 18 \text{ yr}$ of optical photometry from ASAS, KELT, and ASAS-SN we find that TYC 8332-2529-1 has a photometric period of $21.25 \pm 0.04 \text{ d}$. A model of the light curve which attributes the variability to the presence of stars spots explains the light curve shape, changing phase, and results in a reasonable covering fraction. Using spectra

from SALT and LCO we find that the star is in a spectroscopic binary, with a velocity amplitude of 43 km s^{-1} . The X-ray and UV flux that we detect from the position of MKT J170456.2–482100 indicates that this flux and the radio variability may be because the system is an RS CVn type system. There are absorption lines indicating spectral variation within the system, in antiphase to the K-giant spectrum, at a radial velocity amplitude of $12 \pm 10 \text{ km s}^{-1}$. However, we lack sufficient data to identify this with a causative mechanism, as binary companions at this radial velocity do not make physical sense.

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¹¹<https://www.cosmos.esa.int/gaia>, processed by the *Gaia* Data Processing and Analysis Consortium (DPAC, <https://www.cosmos.esa.int/web/gaia/dpac/consortium>)

¹²<http://www.astropy.org>

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APPENDIX A: MKT J170456.2–482100 MEERKAT FLUX DENSITY MEASUREMENTS

Table A1. Flux densities measured by MeerKAT for each epoch for MKT J170456.2–482100. These are the original, unscaled (the primary beam correction of 2.00 ± 0.02 has not been applied) peak flux density measurements from TRAP, (*) indicates measurements that were plotted as upper limits in Figs 2 and 10.

Epoch (MJD)	Epoch (date)	Flux density (mJy)	Uncertainty (mJy)	Local RMS noise (mJy)
58068	2017-11-11	0.048*	0.026	0.036
58222	2018-04-14	0.156	0.018	0.018
58369	2018-09-08	0.043*	0.026	0.039
58375	2018-09-14	0.189	0.025	0.026
58382	2018-09-21	0.116	0.024	0.026
58389	2018-09-28	0.339	0.034	0.035
58396	2018-10-05	0.137	0.027	0.028
58402	2018-10-11	0.227	0.026	0.027
58403	2018-10-12	0.355	0.043	0.044
58410	2018-10-19	0.152	0.028	0.028
58418	2018-10-27	0.098	0.025	0.026
58425	2018-11-03	0.113	0.026	0.027
58432	2018-11-10	0.140	0.026	0.027
58439	2018-11-17	−0.027*	0.040	0.031
58446	2018-11-24	0.067	0.027	0.033
58454	2018-12-02	0.096	0.030	0.031
58460	2018-12-08	0.136	0.029	0.031
58467	2018-12-15	0.198	0.031	0.031
58474	2018-12-22	0.202	0.036	0.036
58481	2018-12-29	0.189	0.035	0.036
58488	2019-01-05	0.192	0.030	0.030
58495	2019-01-12	0.164	0.031	0.032
58502	2019-01-19	−0.005*	0.006	0.028
58509	2019-01-26	0.211	0.029	0.029
58515	2019-02-01	0.025*	0.017	0.030
58523	2019-02-09	0.119	0.033	0.035
58530	2019-02-16	0.077	0.026	0.029
58537	2019-02-23	0.115	0.029	0.031
58543	2019-03-01	0.103	0.027	0.028
58551	2019-03-09	0.091	0.028	0.031
58560	2019-03-18	0.038*	0.023	0.035
58567	2019-03-25	−0.112*	0.408	0.032
58574	2019-04-01	−0.116*	0.373	0.037
58582	2019-04-09	0.108	0.036	0.041
58588	2019-04-15	0.084	0.028	0.031
58593	2019-04-20	0.032*	0.020	0.032
58602	2019-04-29	0.059*	0.029	0.038
58607	2019-05-04	0.033*	0.020	0.031
58614	2019-05-11	0.109	0.032	0.034
58621	2019-05-18	0.097	0.035	0.039
58628	2019-05-25	0.105	0.031	0.033
58634	2019-05-31	0.113	0.034	0.036
58642	2019-06-08	0.015*	0.012	0.040
58650	2019-06-16	0.115	0.032	0.035
58658	2019-06-24	0.051*	0.024	0.031
58664	2019-06-30	2.122*	150.234	0.031
58671	2019-07-07	−0.002*	0.002	0.030
58678	2019-07-14	0.179	0.030	0.030

¹Jodrell Bank Centre for Astrophysics, Department of Physics and Astronomy, The University of Manchester, Manchester M13 9PL, UK

²South African Astronomical Observatory, PO Box 9, Observatory 7935, South Africa

³Southern African Large Telescope, PO Box 9, Observatory 7935, South Africa

⁴Anton Pannekoek Institute, University of Amsterdam, Postbus 94249, NL-1090 GE Amsterdam, the Netherlands

⁵Netherlands Institute for Radio Astronomy (ASTRON), Oude Hoogeveensedijk 4, NL-7991 PD Dwingeloo, the Netherlands

⁶AIM, CEA, CNRS, Université Paris Diderot, Sorbonne Paris Cité, Université Paris-Saclay, F-91191 Gif-sur-Yvette, France

⁷Inter-University Institute for Data Intensive Astronomy, Department of Astronomy, University of Cape Town, Private Bag X3, Rondebosch 7701, South Africa

⁸Department of Physics, Astrophysics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford OX1 3RH, UK

⁹South African Radio Astronomy Observatory, 2 Fir Street, Black River Park, Observatory, Cape Town 7925, South Africa

¹⁰Department of Astrophysics/IMAPP, Radboud University Nijmegen, PO Box 9010, NL-6500 GL Nijmegen, the Netherlands

¹¹Department of Physics and Electronics, Rhodes University, PO Box 94, Grahamstown 6140, South Africa

¹²Racah Institute of Physics, The Hebrew University of Jerusalem, Jerusalem 91904, Israel

¹³Department of Physics, The George Washington University, 725 21st Street NW, Washington, DC 20052, USA

¹⁴Astronomy, Physics and Statistics Institute of Sciences (APSYS), 725 21st Street NW, Washington, DC 20052, USA

¹⁵Department of Astronomy, University of Cape Town, Private Bag X3, Rondebosch 7701, South Africa

¹⁶IAU Office of Astronomy for Development, Cape Town 7935, South Africa

¹⁷International Centre for Radio Astronomy Research – Curtin University, GPO Box U 1987, Perth, WA 6845, Australia

¹⁸Department of Physics, University of Oxford, Keble Road, Oxford OX1 3RH, UK

¹⁹National Radio Astronomy Observatory, Socorro, NM 87801, USA

²⁰Caltech, 1200 E. California Blvd. MC 249-17, Pasadena, CA 91125, USA

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