



HAL
open science

Tip of the red giant branch distance to the nearby dwarf galaxy [TT2009] 25 in the NGC 891 group

Oliver Müller, Rodrigo Ibata, Marina Rejkuba, Lorenzo Posti

► To cite this version:

Oliver Müller, Rodrigo Ibata, Marina Rejkuba, Lorenzo Posti. Tip of the red giant branch distance to the nearby dwarf galaxy [TT2009] 25 in the NGC 891 group. *Astronomy and Astrophysics - A&A*, 2019, 629, pp.L2. 10.1051/0004-6361/201936392 . hal-02400801

HAL Id: hal-02400801

<https://hal.science/hal-02400801v1>

Submitted on 9 Dec 2019

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

LETTER TO THE EDITOR

Tip of the red giant branch distance to the nearby dwarf galaxy [TT2009] 25 in the NGC 891 group

Oliver Müller¹, Rodrigo Ibata¹, Marina Rejkuba², and Lorenzo Posti¹

¹ Observatoire Astronomique de Strasbourg (ObAS), Université de Strasbourg - CNRS, UMR 7550 Strasbourg, France
e-mail: oliver.muller@astro.unistra.fr

² European Southern Observatory, Karl-Schwarzschild Strasse 2, 85748, Garching, Germany

Received tba; accepted tba

ABSTRACT

Dwarf galaxies are key objects for small-scale cosmological tests like the abundance problems or the planes-of-satellites problem. It is therefore a crucial task to get accurate information for as many nearby dwarf galaxies as possible. Using extremely deep, ground-based V and i -band Subaru Suprime Cam photometry with a completeness of $i = 27$ mag, we measure the tip of the red giant branch distance for the dwarf galaxy [TT2009] 25. This dwarf resides in the field around the Milky Way-analog NGC 891. By using a Bayesian approach, we measure a distance of $10.28^{+1.17}_{-1.73}$ Mpc, which is consistent with the distance of NGC 891, thus we confirm it as a member of NGC 891. The dwarf galaxy follows the scaling relations defined by the Local Group dwarfs. We do not find an extended stellar halo around [TT2009] 25. In the small field of view of 100 kpc covered by the survey, only one bright dwarf galaxy and the giant stream are apparent. This is comparable to the Milky Way, where one bright dwarf resides in the same volume, as well as the Sagittarius stream – excluding satellites which are farther away but would be projected in the line-of-sight. It is thus imperative to survey for additional dwarf galaxies in a larger area around NGC 891 to test the abundance of dwarf galaxies and compare it to the number of satellites around the Milky Way.

Key words. Galaxies: distances and redshifts, Galaxies: dwarf, Galaxies: groups: individual: NGC 891, Galaxies: individual: [TT2009] 25

1. Introduction

In the Λ CDM universe structure is built up through hierarchical merging in a bottom-up scenario (e.g. the review by Frenk & White 2012). The smallest building blocks are the dwarf galaxies, which are also the most abundant and most dark matter dominated galaxies in the universe. In recent years, using the dwarf galaxies, several tests of cosmological models have been proposed and conducted, scrutinizing the abundance, the distribution, the motions, and the internal dynamics of the dwarfs (e.g., Moore et al. 1999; Kroupa et al. 2010, 2018; Boylan-Kolchin et al. 2011; Ibata et al. 2014; López-Corredoira & Kroupa 2016; Read et al. 2018; Müller et al. 2018b; van Dokkum et al. 2018; Müller et al. 2019a; Javanmardi et al. 2019; Hagi et al. 2019). For some of these tests the models have successfully found a solution (e.g. Simon & Geha 2007; Sawala et al. 2016; Kroupa et al. 2018; Read et al. 2019), others are still actively discussed in the literature (e.g. Pawlowski et al. 2015; Libeskind et al. 2015; Pawlowski 2018; Müller et al. 2019c; Monelli & Trujillo 2019). To further improve and extend these tests, it is crucial to measure these properties for as many dwarf galaxies in as many different environment as possible (Javanmardi et al. 2016; Geha et al. 2017).

In the nearby universe, one of the most effective ways to estimate the distance of a galaxy is via the tip of the red giant branch (TRGB, Da Costa & Armandroff 1990; Lee et al. 1993). It exploits the maximal brightness a low-mass star can reach along its evolutionary path in the red giant branch (RGB) phase when observed in appropriate optical or near-infrared filters. The observed sharp increase in the (I -band) luminosity function, or cut-

off in the colour-magnitude diagram (CMD), which is theoretically understood (Serenelli et al. 2017) as well as empirically calibrated (Bellazzini 2008), can be successfully used as a standard candle, yielding distance measurements accurate to about 5 to 10 percent (Tully et al. 2015; Müller et al. 2018c; Cohen et al. 2018; Anand et al. 2019; Bennet et al. 2019). In I -band the observed TRGB depends only weakly on age and metallicity, and is typically not affected by the presence of variable stars provided sufficiently well sampled stellar population.

In this work, we present a TRGB distance measurement for the dwarf galaxy [TT2009] 25, discovered by Trentham & Tully (2009) in a wide-field MegaCam based survey of the region around NGC 1023, and more specifically close to the Milky Way analog NGC 891. The Milky Way analog NGC 891 has been targeted to resolve its stellar halo and search for streams through detection of individual RGB stars (Mouhcine et al. 2010) with ground-based Subaru data, accompanied with data from the Hubble space telescope (Rejkuba et al. 2009). While the previous work focused on the stellar halo, which has revealed interaction history with a smaller companion that left streams and loops around NGC 891, in the following we will measure the distance to the dwarf galaxy [TT2009] 25, which resides within the deep Subaru field.

2. Data and photometry

In this work we use data from the study of NGC 891 by Mouhcine et al. (2010) based on deep Subaru Suprime Cam (Miyazaki et al. 2002) imaging. The observations were con-

ducted in the Johnson V -band and the Gunn i -band. A total of 10 hours of good quality data in the V -band and 11 hours in the i -band were obtained. The seeing was better than 0.6 arcsec. This strategy resolved the RGB down to 2 magnitudes below the tip at the distance of 10 Mpc (Mouhcine et al. 2007), corresponding to NGC 891.

The photometry was performed with DAOPHOT (Stetson 1987). For the detailed description of the photometric pipeline we refer to Mouhcine et al. (2010). We have calibrated the photometry with the Pan-STARRS DR2 catalog (Chambers et al. 2016). While calibrating our i -band was a straight forward task, we had to convert Pan-STARRS gr to our V using Lupton (2005). The photometric errors in the i -band range from $i \approx 0.05$ mag at $i \approx 25.8 - 26.0$ mag, i.e., the bright end of the RGB, to $i \approx 0.15$ mag at $i \approx 26.8 - 27.0$ mag (Mouhcine et al. 2010), i.e., approximately a magnitude below the TRGB of NGC 891. This is sufficient to accurately measure the distance at the 5 to 10 percent level. To cull bona-fide stars from background compact galaxies we have applied quality cuts using the PSF fitting sharpness and χ^2 parameters, and rejected objects deviating from the mean photometric errors (see e.g. Fig. 2 of Müller et al. 2019b).

3. TRGB distance measurement

In Fig 1 we present a cut-out of the field around NGC 891, including the dwarf. It is clearly visible and its outskirts are well resolved in stars, while in the centre the crowding gets too high for accurate photometry. In Fig 2 we show the star map of putative RGB stars. To create a color magnitude diagram (CMD), we select all stars within an elliptical annulus, avoiding the innermost stars. This was necessary as blending of individual stars was evident, i.e. their magnitudes were ~ 0.4 mag higher than in the outskirts of the galaxy. We refer to Trujillo et al. (2019) for a recent discussion about the effect of crowding in dwarf galaxies.

We measure the TRGB magnitude using a Bayesian Markov Chain Monte Carlo (MCMC) approach developed by Conn et al. (2011, 2012) and described in detail in Müller et al. (2019b). In this scheme, we approximate the location of RGB stars on the CMD with a power law, and model the contamination of foreground/background objects using a large reference field. The reference field consists of several patches avoiding the stellar loop and the inner halo of NGC 891, with roughly the same distance to NGC 891 as [TT2009] 25, with an area of ≈ 100 square arcmin. A most likely set of m_{TRGB} , a , and c are evaluated, where m_{TRGB} is the TRGB magnitude, a the slope of the power law, and c the contamination factor.

We estimate the distance to [TT2009] 25 by measuring the difference in the distance modulus between NGC 891 and [TT2009] 25. This has the advantage that we can anchor our estimate to the distance of NGC 891 coming from the Hubble space telescope, yielding more accurate distances than what is possible from ground based telescopes, due to the superior image quality. For NGC 891 we use an elliptical annulus placed such that it only covers the outskirts of the galaxy. Using our MCMC scheme, we measure $i_{TRGB} = 26.32^{+0.09}_{-0.09}$ mag. The error bounds are given by the 68% interval of the MCMC chain and already include the photometric uncertainty. This gives a distance of $D = 8.97^{+0.36}_{-0.36}$ Mpc for NGC 891 using the calibration of the TRGB magnitude of -3.44 ± 0.1 (Bellazzini 2008, transformed into AB magnitudes), which indeed differs slightly from the HST distance of $D = 9.73^{+0.94}_{-0.86}$ Mpc (Mouhcine et al. 2007; Rejkuba et al. 2009). For [TT2009] 25 we measure the TRGB to be at $i_{TRGB} = 26.45^{+0.09}_{-0.34}$ mag. In Fig 3 we present the CMD for NGC 891 and the dwarf galaxy [TT2009] 25, as well

as the estimated TRGBs. The difference in the distance modulus between NGC 891 and [TT2009] 25 is 0.12 mag, however, within the uncertainties the two values are consistent. Using a distance of $D = 9.73$ Mpc for NGC 891 (Mouhcine et al. 2007; Rejkuba et al. 2009) this yields the final extinction corrected distance estimate of $D = 10.28^{+1.17}_{-1.73}$ Mpc for [TT2009] 25, with a distance modulus of $(m - M) = 30.06^{+0.23}_{-0.40}$ mag. The uncertainties come from the joint errors of the TRGB detection in the Subaru data ($^{+0.12}_{-0.35}$ mag) and the uncertainty in the TRGB detection in the HST data (± 0.2 mag).

4. Discussion

In the following we discuss the structural properties of [TT2009] 25 and the abundance of dwarf galaxy members around NGC 891.

4.1. Characterizing [TT2009] 25

The distance estimate of [TT2009] 25 is consistent with it being a satellite of NGC 891. This is not surprising, as the velocity $v = 692.0 \pm 58$ km/s of [TT2009] 25 (Karachentsev et al. 2015b) is also consistent with that of NGC 891 ($v = 526 \pm 7$ km/s, Staveley-Smith & Davies 1987). Nonetheless, with the distance estimation we confirm beyond doubt that [TT2009] 25 is a dwarf galaxy associated with NGC 891.

Having the distance information, we can derive its structural parameters. For this task we use Galfit (Peng et al. 2002) – a standard program to measure galaxy profiles. Due to the dwarf being resolved, we used MegaCam data from the CFHT archive, where the dwarf is unresolved. We downloaded the reduced r -band data, calibrated the image using Pan-STARRS DR2 (Magnier et al. 2016), and performed aperture photometry with DAOPHOT to measure the zero point of the image. The dwarf galaxy was then fit with a Sérsic model (Sersic 1968) using Galfit. The derived parameters are shown in Table 1.

The structural parameters follow the relation defined by the dwarfs in the Local Group (Martin et al. 2008; McConnachie 2012) and beyond (e.g. Venhola et al. 2017; Müller et al. 2017, 2018a), and are reminiscent of the dwarf spheroidal galaxy Fornax (Irwin & Hatzidimitriou 1995) in absolute magnitude, effective radius, and ellipticity. The analogy to Fornax goes only so far as [TT2009] 25 has emission in the FUV in GALEX (Martin et al. 2005), indicating star formation. So [TT2009] 25 is rather a transition type dwarf than a dwarf spheroidal. Furthermore, no atomic hydrogen is clearly detected with deep radio interferometry at the position of this dwarf galaxy (Oosterloo et al. 2007), which is expected for bound dwarf galaxies around giants (Spekkens et al. 2014).

The stellar profile of [TT2009] 25 seems to have a rather sharp cut-off (visible in Fig 2). This argues against an extended stellar halo around this dwarf. The existence or absence of such extended stellar halos has been discussed in the literature (e.g., Minniti & Zijlstra 1996; Rejkuba et al. 2000; Minniti et al. 2003; Gallart et al. 2004; Hidalgo et al. 2009; Greggio et al. 2018), with growing evidence that even low mass galaxies may have experienced hierarchical growth (e.g. Annibali et al. 2016) and could thus host extended stellar envelopes. However, the superposition of the outer halo population of NGC 891 in front of [TT2009] 25 would require deeper data to detect any such low surface brightness envelope around the dwarf galaxy (if present).

As a final note, [TT2009] 25 seems not to be associated with the giant stream around NGC 891. This is evident when studying

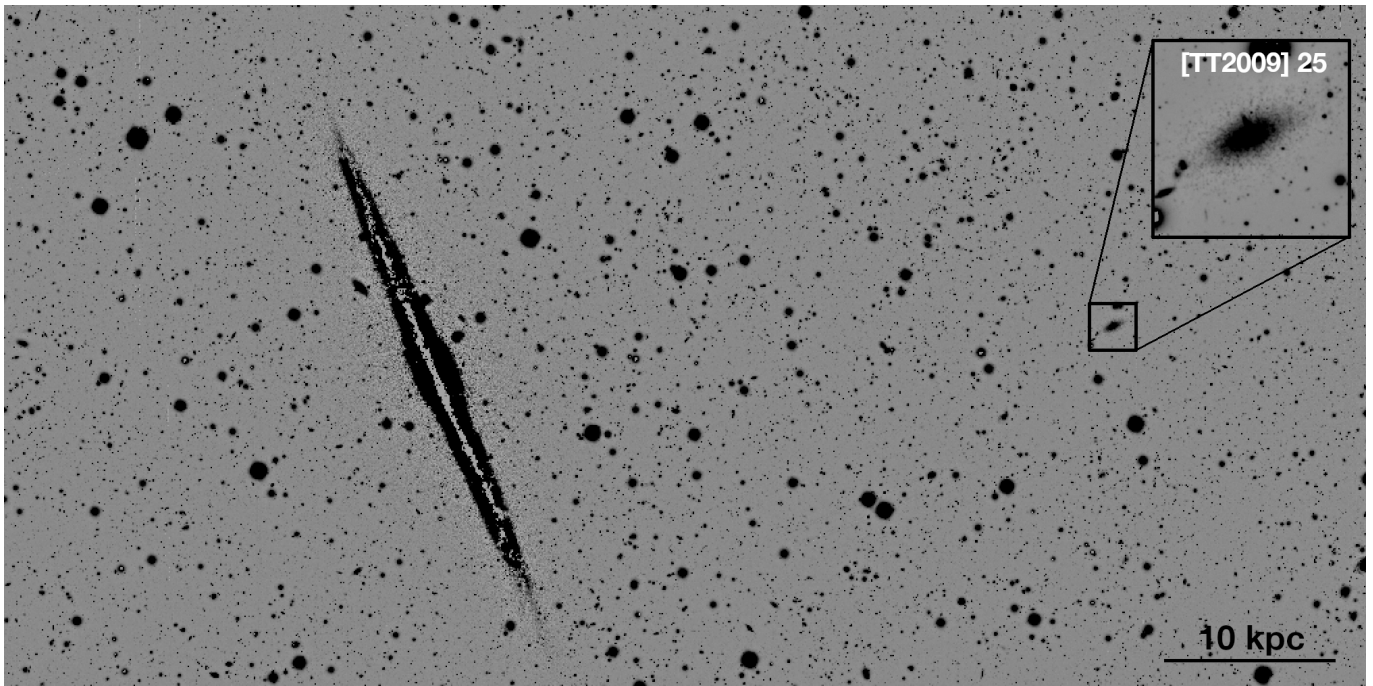


Fig. 1. The field around NGC 891 from our deep, stacked *i*-band Subaru Suprime Cam image. The box highlighting the dwarf galaxy [TT2009] 25 has a side of 1 arcmin. The scale indicates 10 kpc at 10 Mpc.

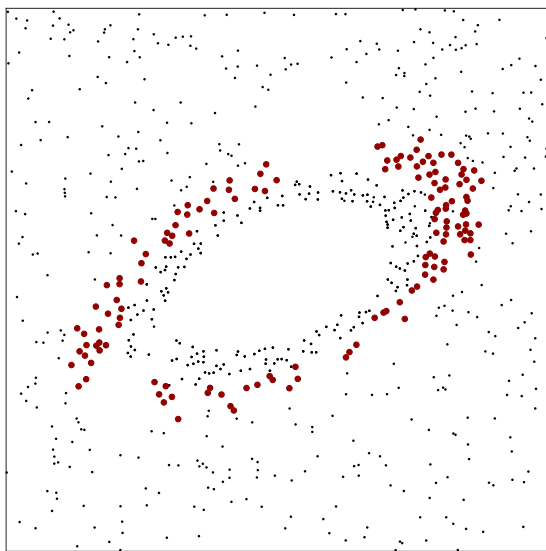


Fig. 2. Stars map around [TT2009] 25. Indicated are objects fainter than $i = 25.5$ mag. The large red dots correspond to the stars selected for the color magnitude diagram. The inner part of the galaxy is obfuscated due to the crowding of stars. One side is 100 arcsec.

Fig. 1 of Mouhcine et al. (2010), where the RGB stars in the full field are displayed and [TT2009] 25 is well visible as separated overdensity of stars. The dwarf galaxy seems to be well away from the stream and ~ 500 kpc behind the halo of NGC 891.

4.2. Abundance of dwarf galaxies around NGC 891

Apart from the giant stream surrounding NGC 891, [TT2009] 25 is the only resolved object in our field. This is not unexpected, as the field of view only covers ~ 100 kpc at the distance of NGC 891. The spiral galaxy NGC 891 is thought to be a Milky

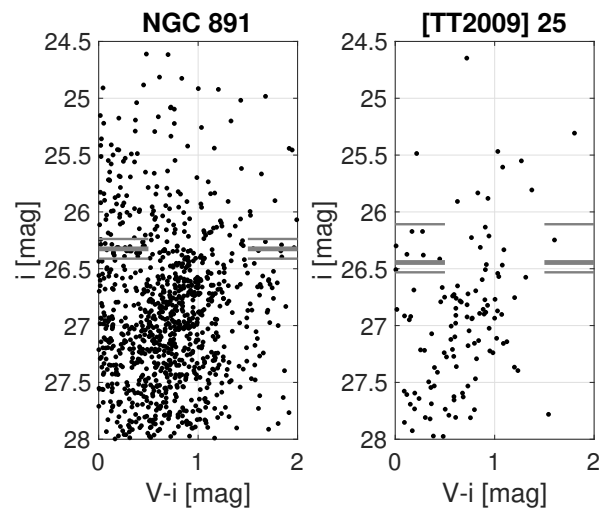


Fig. 3. The color magnitude diagram of NGC 891 (left) and [TT2009] 25 (right). The gray lines denote the best estimated TRGB magnitude and the corresponding uncertainties.

Way analog in size and mass. When we take the dwarf galaxy satellites around the Milky Way within a sphere of 50 kpc and brighter than -10 mag (McConnachie 2012; Pawlowski et al. 2013), we find only the Sagittarius dSph and the Large Magellanic Cloud, i.e. two satellites. This is quite comparable to what we find for NGC 891: a stream that is reminiscent of Sagittarius and a moderately bright dwarf galaxy. At fainter magnitudes though, the Milky Way hosts at least 10 dwarf galaxies (McConnachie 2012), with the numbers still increasing. It is thus to be expected that many low mass and low surface brightness satellites await discovery within the observed field around NGC 891. Indeed, Ibata et al. (2009) identified 6 overdensities within the much smaller HST/ACS area imaged close to the main galaxy disk, that could be either candidate satellites or satellite

Table 1. Properties of [TT2009] 25.

	[TT2009] 25
RA (J2000)	02:21:12.2
DEC (J2000)	+42:21:50
$(m - M)_0$ (mag)	$30.06^{+0.23}_{-0.40}$
Distance (Mpc)	$10.28^{+1.17}_{-1.73}$
A_i (mag)	0.09
m_r (mag)	16.88 ± 0.06
M_r (mag)	$-13.18^{+0.24}_{-0.40}$
L_r ($10^6 M_\odot$)	$16.0^{+7.1}_{-3.2}$
$r_{eff,r}$ (arcsec)	13.02 ± 0.04
$r_{eff,r}$ (pc)	649^{+76}_{-111}
$\mu_{eff,r}$ (mag arcsec $^{-2}$)	24.45 ± 0.06
PA (north to east)	116.9 ± 0.1
e ($1 - b/a$)	0.52 ± 0.01
Sérsic index n	1.01 ± 0.01

debris. In a wider field of view, Karachentsev et al. (2015a) found no new dwarf galaxy candidate based on a 12 hour exposure with an amateur telescope, apart of [TT2009] 25 and [TT2009] 30. This does not mean that NGC 891 is sparsely populated. Within 300 kpc, the Local Volume catalog (Karachentsev et al. 2004, 2013) yields four more suspected irregular type dwarf galaxies based on their radial velocities: DDO 22, DDO 24, UGC 1807 and UGC 2172.

5. Summary and conclusion

Using extremely deep V and i band images taken with the Subaru Suprime Cam we have resolved the upper part of the red giant branch of the dwarf galaxy [TT2009] 25. We have measured the tip of the red giant branch brightness using a Markov Chain Monte Carlo method and have derived an accurate distance for this dwarf galaxy, which is consistent with the distance of NGC 891. We thus confirm its membership.

Using data from the CFHT archive, we have measured the structural parameters of [TT2009] 25, which are compatible in absolute magnitude, effective radius, and ellipticity to other known dwarf galaxies. Most notably, it is reminiscent of the Local Group dwarf spheroidal galaxy Fornax in morphology, even though the detection of FUV emission makes [TT2009] 25 rather a transitional type dwarf. Still, this makes the dwarf galaxy an interesting target for follow-up studies. For example, the globular cluster distribution around Fornax yields strong constraints about the shape of its dark matter halo, favoring a core-like profile over a cusp (Cole et al. 2012; Boldrini et al. 2019; Orkney et al. 2019). Identifying globular clusters around [TT2009] 25 could therefore bring new insights into the cusp/core problem (de Blok 2010), as well as directly measuring the velocity dispersion profile, which should be possible with future facilities like the Thirty Meter Telescope.

Due to the limited field of view, we only cover a region of 50 kpc around NGC 891. In this field, we find one resolved dwarf galaxy. This is comparable to what is expected from the satellite distribution around the Milky Way, where two bright/classical dwarf galaxies reside within a radius of 50 kpc. However, this does not include dwarf galaxies which can be farther away from the host, but projected along our line-of-sight, which would

increase the number of expected satellites observed around NGC 891. It is thus imperative to search for dwarf galaxies in a larger area around NGC 891 to test the abundance of dwarf galaxies and compare it to the Milky Way.

Acknowledgements. We thank the referee for the constructive report, which helped to clarify and improve the manuscript. O.M. is grateful to the Swiss National Science Foundation for financial support. The authors thank Filippo Fraternali for discussions about the HI content in the field, and Michele Bellazzini for discussion about the RGB tip calibration in different filter systems.

References

- Anand, G. S., Tully, R. B., Rizzi, L., Shaya, E. J., & Karachentsev, I. D. 2019, arXiv e-prints, arXiv:1905.11416
- Annibali, F., Nipoti, C., Ciotti, L., et al. 2016, ApJ, 826, L27
- Bellazzini, M. 2008, Mem. Soc. Astron. Italiana, 79, 440
- Bennet, P., Sand, D. J., Crnojević, D., et al. 2019, arXiv e-prints, arXiv:1906.03230
- Boldrini, P., Mohayaee, R., & Silk, J. 2019, MNRAS, 485, 2546
- Boylan-Kolchin, M., Bullock, J. S., & Kaplinghat, M. 2011, MNRAS, 415, L40
- Chambers, K. C., Magnier, E. A., Metcalfe, N., et al. 2016, arXiv e-prints, arXiv:1612.05560
- Cohen, Y., van Dokkum, P., Danieli, S., et al. 2018, ApJ, 868, 96
- Cole, D. R., Dehnen, W., Read, J. I., & Wilkinson, M. I. 2012, MNRAS, 426, 601
- Conn, A. R., Ibata, R. A., Lewis, G. F., et al. 2012, ApJ, 758, 11
- Conn, A. R., Lewis, G. F., Ibata, R. A., et al. 2011, ApJ, 740, 69
- Da Costa, G. S. & Armandroff, T. E. 1990, AJ, 100, 162
- de Blok, W. J. G. 2010, Advances in Astronomy, 2010, 789293
- Frenk, C. S. & White, S. D. M. 2012, Annalen der Physik, 524, 507
- Gallart, C., Stetson, P. B., Hardy, E., Pont, F., & Zinn, R. 2004, ApJ, 614, L109
- Geha, M., Wechsler, R. H., Mao, Y.-Y., et al. 2017, ApJ, 847, 4
- Greggio, L., Falomo, R., & Scarpa, R. 2018, ApJ, 861, 81
- Haghi, H., Kroupa, P., Banik, I., et al. 2019, MNRAS, 487, 2441
- Hidalgo, S. L., Aparicio, A., Martínez-Delgado, D., & Gallart, C. 2009, ApJ, 705, 704
- Ibata, N. G., Ibata, R. A., Famaey, B., & Lewis, G. F. 2014, Nature, 511, 563
- Ibata, R., Mouhcine, M., & Rejkuba, M. 2009, MNRAS, 395, 126
- Irwin, M. & Hatzidimitriou, D. 1995, MNRAS, 277, 1354
- Javanmardi, B., Martínez-Delgado, D., Kroupa, P., et al. 2016, A&A, 588, A89
- Javanmardi, B., Raouf, M., Khosroshahi, H. G., et al. 2019, ApJ, 870, 50
- Karachentsev, I. D., Karachentseva, V. E., Huchtmeier, W. K., & Makarov, D. I. 2004, AJ, 127, 2031
- Karachentsev, I. D., Makarov, D. I., & Kaisina, E. I. 2013, AJ, 145, 101
- Karachentsev, I. D., Riepe, P., Zilch, T., et al. 2015a, Astrophysical Bulletin, 70, 379
- Karachentsev, I. D., Sharina, M. E., Makarov, D. I., Perepelitsyna, Y. A., & Safonova, E. S. 2015b, Astrophysics, 58, 309
- Kroupa, P., Banik, I., Haghi, H., et al. 2018, Nature Astronomy, 2, 925
- Kroupa, P., Famaey, B., de Boer, K. S., et al. 2010, A&A, 523, A32
- Kroupa, P., Haghi, H., Javanmardi, B., et al. 2018, Nature, 561, E4
- Lee, M. G., Freedman, W. L., & Madore, B. F. 1993, ApJ, 417, 553
- Libeskind, N. I., Hoffman, Y., Tully, R. B., et al. 2015, MNRAS, 452, 1052
- López-Corredoira, M. & Kroupa, P. 2016, ApJ, 817, 75
- Lupton, R. 2005, Transformations between SDSS magnitudes and other systems <https://www.sdss3.org/dr10/algorithms/sdssUBVRITransform.php>
- Magnier, E. A., Schlafly, E. F., Finkbeiner, D. P., et al. 2016, arXiv e-prints [arXiv:1612.05242]
- Martin, D. C., Fanson, J., Schiminovich, D., et al. 2005, ApJ, 619, L1
- Martin, N. F., de Jong, J. T. A., & Rix, H.-W. 2008, ApJ, 684, 1075
- McConnachie, A. W. 2012, AJ, 144, 4
- Minniti, D., Borissova, J., Rejkuba, M., et al. 2003, Science, 301, 1508
- Minniti, D. & Zijlstra, A. A. 1996, ApJ, 467, L13
- Miyazaki, S., Komiyama, Y., Sekiguchi, M., et al. 2002, PASJ, 54, 833
- Monelli, M. & Trujillo, I. 2019, arXiv e-prints, arXiv:1907.03761
- Moore, B., Ghigna, S., Governato, F., et al. 1999, ApJ, 524, L19
- Mouhcine, M., Ibata, R., & Rejkuba, M. 2010, ApJ, 714, L12
- Mouhcine, M., Rejkuba, M., & Ibata, R. 2007, MNRAS, 381, 873
- Müller, O., Famaey, B., & Zhao, H. 2019a, A&A, 623, A36
- Müller, O., Jerjen, H., & Binggeli, B. 2017, A&A, 597, A7
- Müller, O., Jerjen, H., & Binggeli, B. 2018a, A&A, 615, A105
- Müller, O., Pawlowski, M. S., Jerjen, H., & Lelli, F. 2018b, Science, 359, 534
- Müller, O., Rejkuba, M., & Jerjen, H. 2018c, A&A, 615, A96
- Müller, O., Rejkuba, M., Pawlowski, M. S., et al. 2019b, arXiv e-prints, arXiv:1907.02012

- Müller, O., Rich, R. M., Román, J., et al. 2019c, arXiv e-prints [arXiv:1903.07285]
- Oosterloo, T., Fraternali, F., & Sancisi, R. 2007, *AJ*, 134, 1019
- Orkney, M. D. A., Read, J. I., Petts, J. A., & Gieles, M. 2019, *MNRAS*, 1624
- Pawlowski, M. S. 2018, *Modern Physics Letters A*, 33, 1830004
- Pawlowski, M. S., Famaey, B., Merritt, D., & Kroupa, P. 2015, *ApJ*, 815, 19
- Pawlowski, M. S., Kroupa, P., & Jerjen, H. 2013, *MNRAS*, 435, 1928
- Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H.-W. 2002, *AJ*, 124, 266
- Read, J. I., Walker, M. G., & Steger, P. 2018, *MNRAS*, 481, 860
- Read, J. I., Walker, M. G., & Steger, P. 2019, *MNRAS*, 484, 1401
- Rejkuba, M., Minniti, D., Gregg, M. D., et al. 2000, *AJ*, 120, 801
- Rejkuba, M., Mouhcine, M., & Ibata, R. 2009, *MNRAS*, 396, 1231
- Sawala, T., Frenk, C. S., Fattahi, A., et al. 2016, *MNRAS*, 457, 1931
- Serenelli, A., Weiss, A., Cassisi, S., Salaris, M., & Pietrinferni, A. 2017, *A&A*, 606, A33
- Sersic, J. L. 1968, *Atlas de galaxias australes*
- Simon, J. D. & Geha, M. 2007, *ApJ*, 670, 313
- Spekkens, K., Urbancic, N., Mason, B. S., Willman, B., & Aguirre, J. E. 2014, *ApJ*, 795, L5
- Staveley-Smith, L. & Davies, R. D. 1987, *MNRAS*, 224, 953
- Stetson, P. B. 1987, *PASP*, 99, 191
- Trentham, N. & Tully, R. B. 2009, *MNRAS*, 398, 722
- Trujillo, I., Beasley, M. A., Borlaff, A., et al. 2019, *MNRAS*, 486, 1192
- Tully, R. B., Libeskind, N. I., Karachentsev, I. D., et al. 2015, *ApJ*, 802, L25
- van Dokkum, P., Danieli, S., Cohen, Y., et al. 2018, *Nature*, 555, 629
- Venhola, A., Peletier, R., Laurikainen, E., et al. 2017, *A&A*, 608, A142