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The CFHT large area U -band deep survey (CLAUDS)

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ABSTRACT

The Canada–France–Hawaii Telescope (CFHT) Large Area U -band Deep Survey (CLAUDS) uses data taken with the MegaCam mosaic imager on CFHT to produce images of 18.60 deg^2 with median seeing of $\text{FWHM} = 0.92 \text{ arcsec}$ and to a median depth of $U = 27.1 \text{ AB}$ (5σ in 2 arcsec apertures), with selected areas that total 1.36 deg^2 reaching a median depth of $U = 27.7 \text{ AB}$. These are the deepest U -band images assembled to date over this large an area. These data are located in four fields also imaged to comparably faint levels in *grizy* and several narrowband filters as part of the Hyper Suprime–Cam (HSC) Subaru Strategic Program (HSC-SSP). These CFHT and Subaru data sets will remain unmatched in their combination of area and depth until the advent of the Large Synoptic Survey Telescope. This paper provides an overview of the scientific motivation for CLAUDS and gives details of the observing strategy, observations, data reduction, and data merging with the HSC-SSP. Three early applications of these deep data are used to illustrate the potential of the data set: deep U -band galaxy number counts, $z \sim 3$ Lyman break galaxy selection, and photometric redshifts improved by adding CLAUDS U to the Subaru HSC *grizy* photometry.

Key words: galaxies: formation – galaxies: haloes – galaxies: statistics – dark matter – large-scale structure of universe – cosmology: observations.

1 INTRODUCTION

A major achievement of observational cosmology has been the detailed characterization of large-scale structure (LSS) traced by galaxies in both the local and high- z Universe, and its successful explanation in the Λ CDM cosmological framework. However, within this broad Λ CDM paradigm, our understanding of the formation of structures on the scales of galaxies contains vast gaps: we still do not fully understand the link between galaxy

properties and their host dark matter haloes or the baryonic physics from gaseous infall, through star formation (SF), to SF-regulating feedback processes. We do not yet fully know how the galaxies around us today evolved from the objects we see at high redshifts.

A powerful approach to tackle such issues, spectacularly demonstrated at low redshift by the Sloan Digital Sky Survey (SDSS; York et al. 2000), is with deep, large-area imaging and spectroscopic surveys. Such surveys make it possible to derive galaxy properties – such as stellar mass (M_*), star formation rate (SFR), and morphology – for galaxies drawn from large, representative volumes, and then to use statistical tools to link these properties to those of their host DM haloes and to understand the role played by local and large-scale environment.

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However, because galaxies evolve over billions of years, truly large-volume surveys (even the $\sim 2 \text{ deg}^2$ COSMOS field; Scoville et al. 2007) is affected by cosmic variance – see e.g. Arcila-Osejo & Sawicki 2013) need to be extended beyond the local Universe to span the vast lookback times over which this evolution takes place. This extension to high redshift is one of the key motivations for the Hyper Suprime-Cam Subaru Strategic Program (HSC-SSP; Aihara et al. 2018a) currently underway on the Subaru telescope using the Hyper Suprime-Cam imager (HSC; Miyazaki et al. 2018), as well as for its planned follow-up with the Prime Focus Spectrograph (PFS; Takada et al. 2014).

Of particular interest here are the Deep and UltraDeep components of the HSC-SSP, which image the sky in five broadband (*grizy*) and four narrow-band (NB) filters ($\lambda = 387, 816, 921, 1010 \text{ nm}$) to unprecedented combinations of depth and area: when completed, the Deep component will reach $i_{\text{lim}} \sim 27.1 \text{ AB}$ over $\sim 26 \text{ deg}^2$ (5σ limits in 2 arcsec apertures) and UltraDeep will reach $i_{\text{lim}} \sim 27.7 \text{ AB}$ over 3.5 deg^2 (see Aihara et al. 2019). The four Deep fields (E-COSMOS, XMM-LSS, ELAIS-N1, DEEP2-3), along with the two UltraDeep fields (COSMOS and SXDS, which are embedded within two of the Deep fields, namely E-COSMOS and XMM-LSS, respectively) are in well-studied areas of the sky rich in ancillary data, including extensive spectroscopy and IR and X-ray imaging. Together, these data are ideal for galaxy evolution studies as they provide an unprecedented combination of depth and area that is key for assembling large samples spanning a range of properties in a variety of environments, dominating over cosmic variance, beating down statistical noise, and finding rare objects. However, these areas lack *U*-band data of comparable depth and areal coverage.

While we can learn much from data at longer wavelengths, *U*-band observations are critical for several areas of research, including for studies at $z \lesssim 0.7$ and $z \sim 2-3$, as we discuss in this paper. A number of deep and wide *U*-band imaging surveys have been undertaken in the past (see Fig. 1), including some that overlap parts of the HSC-SSP Deep and UltraDeep fields, but none of them have the combination of depth and area needed to match the HSC-SSP Deep and UltraDeep *grizy* + NB observations. *U*-band data of depth comparable to the HSC-SSP images are thus essential, and with this in mind we carried out the CFHT Large Area *U*-band Deep Survey (CLAUDS). This survey was enabled by the MegaCam imager (Boulade et al. 2003) which, in contrast to HSC and most other imagers on large-aperture telescopes, is unique in its combination of areal coverage ($\sim 1 \text{ deg}^2$) and *U*-band sensitivity. The stacked CLAUDS images that we produced reach a median depth of $U = 27.1 \text{ AB}$ (5σ in 2 arcsec apertures) over 18.60 deg^2 in the HSC-SSP Deep layer, with selected areas within the UltraDeep regions that total 1.36 deg^2 reaching a median depth of $U = 27.7 \text{ AB}$.

As a standalone survey CLAUDS is thus unmatched in the space of *U*-band surveys (see Fig. 1). Combined, CLAUDS+HSC-SSP will be unsurpassed in their combination of depth, area, and wavelength coverage until the advent of the Large Synoptic Survey Telescope (LSST; Abell et al. 2009; Ivezić et al. 2019). Together, these data will allow an unprecedented exploration of cosmic evolution at $0 \lesssim z \lesssim 3$. Even in the era of highly multiplexed spectroscopy, photometric redshifts (photo- z) are essential for large and deep galaxy samples. At intermediate redshifts, photo- z require *U*-band photometry to bracket the Balmer and 4000 \AA breaks (e.g. Connolly et al. 1995; Sawicki, Lin & Yee 1997; Sorba & Sawicki 2011; see also Section 4.2 of this paper). Moreover, *U*-band is important for constraining star formation rates (SFRs) of galaxies

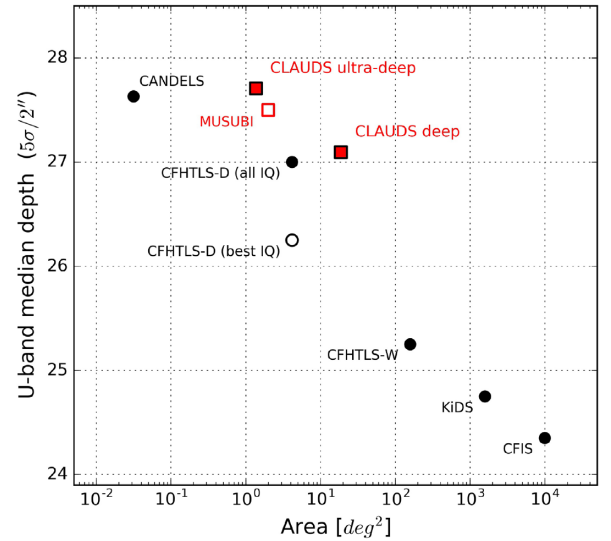


Figure 1. CLAUDS median depths are shown in the context of recent and ongoing *U*-band surveys: Cosmic Assembly Near-IR Deep Extragalactic Legacy Survey (CANDELS; Grogin et al. 2011); the Wide and Deep components of the CFHT Legacy Survey (CFHTLS; Hudelot et al. 2012) – CFHTLS-W and CFHTLS-D, respectively; the Kilo Degree Survey (KiDS; de Jong et al. 2017); and the Canada-France Imaging Survey (CFIS, Ibata et al. 2017). The red open square is for the MegaCam Ultra-deep Survey with *U*-Band Imaging (MUSUBI; W.-H. Wang et al., in preparation) whose data are incorporated into the CLAUDS stacks. Note that CFHTLS-D and MUSUBI contain CFHT u^* data in the COSMOS and SXDS fields that we incorporate into our CLAUDS stacks. The CLAUDS ultra-deep area point includes both u and u^* stacks in the central COSMOS field in addition to the u^* data in the SXDS field that forms part of our XMM-LSS field.

at these redshifts, including through spectral energy distribution (SED) fitting (e.g. Sawicki & Yee 1998; Salim et al. 2009; Sawicki 2012a). The photo- z precision achievable with the CLAUDS+HSC-SSP data will allow rudimentary measures of environment (e.g. Malavasi et al. 2016; Moutard et al. 2018), while the images are sufficient for basic morphological measurements [CLAUDS *U*-band data has median seeing of 0.92 arcsec (Section 3.2.2), and HSC *i*-band has 0.62 arcsec (Aihara et al. 2019)], and identification of galaxy interactions using morphological features for which very deep images are key (Bridge, Carlberg & Sullivan 2010).

The volumes probed by the CLAUDS+HSC-SSP data (e.g. 2×10^7 comoving Mpc^3 in the $z = 0.6-0.8$ slice) approach those spanned by the SDSS main galaxy sample (Strauss et al. 2002) and so contain galaxy samples that are large enough to be split by environment, SFR, stellar mass, and morphology. They are also large enough to overcome cosmic variance and will thus provide definitive measurements of the galaxy UV luminosity function (UVLF) and its evolution out to $z \sim 3$ via both photometric redshifts and, at $z \sim 2-3$, with BM/BX/*U*-drop Lyman Break Galaxy (LBG) techniques (e.g. Steidel et al. 2004; Sawicki & Thompson 2006a; Hildebrandt et al. 2009a; Sawicki 2012b). Using photometric redshifts or drop-out techniques, these data will also allow definitive measurements of galaxy clustering (e.g. Adelberger & Steidel 2005; Savoy et al. 2011) and other methods of linking galaxies and their dark matter haloes, such as weak lensing and halo occupation distribution (HOD) modelling (e.g. Ouchi et al. 2004; Foucaud et al. 2010; Coupon et al. 2012). Using $z \sim 2-3$ BM/BX/LBG samples will permit magnification bias measurements of cluster masses at $z > 1$, where traditional lensing is inadequate (Broadhurst, Taylor &

Table 1. CLAUDS U filters. Filter name is the name adopted by CLAUDS and follows the naming convention recommended by CFHT. Alternate name is the name used in some CFHT documentation; CLAUDS does not use these alternate names, except for uS which we use to mean u^* in ASCII file names and related data products. CFHT ID is the number the filter is identified with in the CFHT filter data base. Filter parameters are taken from the MegaCam filter data base (<https://www.cfht.hawaii.edu/Instruments/Filters/megaprime.html>).

Filter name	Alternate name	CFHT ID	Central wavelength (Å)	Bandwidth (Å)
u	u'	9302	3538	868
u^*	uS, U'	9301	3743	758

Peacock 1995; Hildebrandt, van Waerbeke & Erben 2009b; Tudorica et al. 2017). The U -band is also vital for providing the continuum measurement for line-emitters located in HSC’s NB387 filter and creating large samples of $z = 2.2$ Lyman α emitters (LAEs).

The CLAUDS+HSC-SSP $z \sim 2-3$ BM/BX/LBG and LAE samples will serve as low-redshift reference for higher- z studies that use large areas and find rare, intrinsically luminous objects (e.g. Konno et al. 2018; Matsuoka et al. 2018; Ono et al. 2018). The large areas covered by these data are also needed to find large samples of other intrinsically rare objects, including proto-clusters at $z \sim 2-3$ (e.g. Toshikawa et al. 2016), $z \sim 3$ Lyman Continuum Emitters (e.g. Iwata et al. 2009), $z \sim 3$ quasars (e.g. Fontanot et al. 2007), Galactic white dwarfs and – in combination with the NB387 filter – several hundred Lyman α Blobs (LABs; e.g. Steidel et al. 2000; Matsuda et al. 2011). Indeed, some such rare objects have already been found with these data, including Lyman continuum-emitting galaxies (Bassett et al. 2019) and AGN (I. Iwata et al., in preparation) at $z \sim 3$, and low-mass AGN at $z < 1$ (G. Halevi et al., submitted; G. Halevi et al., in preparation).

This paper focuses on giving an overview of the CLAUDS project. It contains a description of the CLAUDS fields and observations (Section 2), data reductions and an assessment of data quality (Section 3), and several illustrative uses of the data (Section 4): U -band galaxy number counts (Section 4.1), photometric redshifts (Section 4.2), and U -band drop-out selection (Section 4.3). We use AB magnitudes throughout and assume the $(\Omega_M, \Omega_\Lambda, H_0) = (0.3, 0.7, 70 \text{ km s}^{-1} \text{ Mpc}^{-1})$ cosmology. We use u^* and u to refer to the two specific MegaCam filters used in this project and U to mean the general broad-band wavelength region $\sim 3000-4000 \text{ \AA}$ or when we want to refer to the two filters, u^* and/or u , without specifying either one in particular. Table 1 gives some key details of the two U filters (u and u^*) used by CLAUDS, while their transmission curves are shown in Fig. 2.

2 CLAUDS FIELDS AND OBSERVATIONS

2.1 Instrument and filters

The instrument used for CLAUDS observations is MegaCam (Boulade et al. 2003), a wide-field optical imager mounted on MegaPrime, the prime focus on the 3.6m CFHT. While now surpassed in many respects by newer wide-field imagers such as HSC (Miyazaki et al. 2018) and DECam (Flaughner et al. 2015), MegaCam remains unrivalled in its sensitivity in the blue part of the spectrum. To take advantage of the excellent natural seeing at the CFHT site, MegaPrime includes a wide-field corrector that delivers uniform image quality over the entire field of view and an image stabilizing unit that effectively removes telescope wind shake.

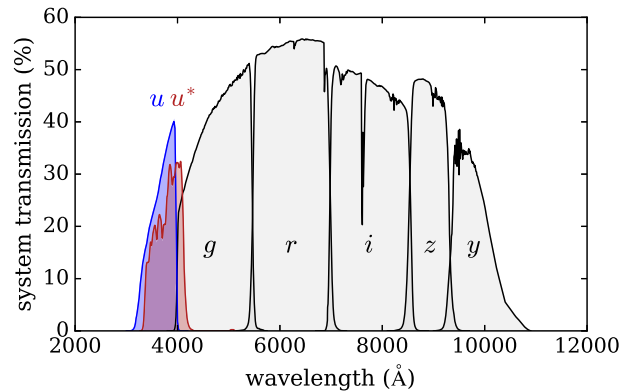


Figure 2. System transmission curves for the CLAUDS u and u^* as well as the HSC $grizy$ bandpasses. The filter transmission curves are taken from the CFHT website and also include the effects of telescope/instrument optics, CCD response, and 1.25 airmasses of typical Mauna Kea atmospheric extinction as described in Betoule et al. (2013); the HSC filter transmission curves are from Kawanomoto et al. (2018) and also include telescope/instrument optics, CCD response, and 1.2 airmasses of extinction, all provided to us by the HSC-SSP collaboration. Note that the new u filter (shown in blue) is significantly different from the old u^* (red): in addition to a higher throughput, its cut-on, central, and cut-off wavelengths are all bluer than those for the old u^* . The new u filter also does not suffer from the old u^* filter’s red leak at $\sim 5000 \text{ \AA}$.

Furthermore, just before the start of the CLAUDS observations, the CFHT enclosure was retrofitted with venting louvres that increase airflow through the dome (Bauman et al. 2014) and are thought to improve seeing by ~ 0.1 arcsec compared to previous CFHT MegaCam imaging programs such as the Canada-France Hawaii Telescope Legacy Survey (CFHTLS; Hudelot et al. 2012) and the Next Generation Virgo Survey (NGVS; Ferrarese et al. 2012).

MegaCam consists of 40 back-illuminated CCDs with good quantum efficiency in the near-UV, fabricated by e2v Technologies. Individual CCDs measure 2048×4612 pixels; the pixels are $13.5 \mu\text{m} \times 13.5 \mu\text{m}$ in size projecting to $0.187 \text{ arcsec} \times 0.187 \text{ arcsec}$ on the sky. The 40 CCDs are arranged in a 4×11 mosaic with the four corner positions empty. The spacing between CCDs is approximately 13 arcsec, with larger gaps (~ 80 arcsec) between the uppermost and lowermost rows and the other CCDs. This configuration spans $1.21 \text{ deg} \times 0.98 \text{ deg}$, including the inter-chip gaps. A single exposure with all 40 CCDs produces a 378 megapixel image of 1.02 deg^2 (exposed sky area). The CCDs take 40 s to read out when – as is standard – two amplifiers are used per chip. The typical read noise is $\sim 5 e^-/\text{pixel}$. The cosmetic quality of the detectors is excellent and only ~ 0.2 percent of the pixels is unresponsive.

MegaCam is equipped with a filter jukebox that can accommodate up to eight filters. However, the original MegaCam filters, $u^*g'r'i'z'$, used in previous programs such as CFHTLS, were only large enough to cover the central 9×4 block of CCDs, thus illuminating 36 CCDs but leaving the outer four CCDs un-illuminated. In late 2014 the filter set was upgraded with the purchase of new $ugriz$ filters. These new filters have significantly better throughputs than the original filters. The new filters are also physically larger, allowing all 40 CCDs of the array to be illuminated, instead of the 36 that were accessible by the old filters, thereby giving an ~ 11 percent improvement in collecting area.

Of particular interest to our program are the U -band filters: CLAUDS uses both the old u^* filter and the new u filter (see

Fig. 2 for their filter transmission curves and Table 1 for some key details). CLAUDS consists of U -band data in four separate $\sim 5 \text{ deg}^2$ fields: E-COSMOS, ELAIS-N1, DEEP2-3, and XMM-LSS. In two of these fields (ELAIS-N1 and DEEP2-3) CLAUDS data consist of u images only; one field (XMM-LSS) contains u^* only; and one (E-COSMOS) has both u and u^* . The two filters, u^* and u are sufficiently different that we treat them entirely separately, producing separate u and u^* stacked images and object catalogues even in the $\sim 1 \text{ deg}^2$ where we have overlap in u^* and u imaging.

2.2 The CLAUDS fields

CLAUDS consists of four fields (the E-COSMOS, ELAIS-N1, DEEP2-3, and XMM-LSS fields), each covering $\sim 4\text{--}6 \text{ deg}^2$ (Fig. 3) with several abutting or partially overlapping MegaCam subfields. The field locations (Table 2) and layouts (Fig. 3) were chosen to maximally overlap with the *grizy* + NB imaging of the Deep Layer of the HSC-SSP (Aihara et al. 2018a) while respecting the on-the-sky footprints of both the HSC and MegaCam and, in some cases, the existence of archival data. In the ELAIS-N1 and DEEP2-3 fields all U -band data were taken by us with the new u filter.

In the XMM-LSS field significant amounts of high-quality data taken with the old u^* filter were available in the CFHT archive or – while still proprietary at the time – were made available to us by the Mega-Cam Survey with U -band Imaging (MUSUBI; W.-H. Wang et al., in preparation). To benefit from these data, we used the old u^* filter to observe in this field. In E-COSMOS, ELAIS-N1, and DEEP2-3 we used the u filter, taking advantage of its better throughput and area. However, the central square degree of our E-COSMOS field, which corresponds to the intensively studied COSMOS field (?), also has significant archival and MUSUBI u^* -band data. We reprocessed these u^* data through our pipeline to produce a very deep u^* stack. Because we also observed this central area of E-COSMOS with the u filter, we have $\sim 1 \text{ deg}^2$ of the sky with very deep data in both u and u^* . This overlap allows us to compare the u and u^* data directly; it also increases the effective depth of U -band imaging available in this important and well-studied field.

2.3 Observations

CLAUDS is a program that combines time from three CFHT partners, namely Canada, France, and China. Observing was done in queue mode over five semesters (2014B–2016B) with typical allocations of 80–86 h per semester with the exception of the final, fifth semester in which a 40 h allocation was used to finish the remaining observations. In total, the CLAUDS project was allocated 376 h from the three Agencies; CFHT experience over many years shows that a MegaCam observing night results in 5.5 h¹ of observing time on average (accounting for weather and telescope and instrument problems and averaging over the seasons and years), so these 376 allocated hours correspond to an allocation of ~ 68 classical-mode observing nights. Moreover, because CLAUDS observing was usually allocated very high ranking in the observing queue, the data were taken in significantly better conditions (seeing, transparency) than would have happened in a random set of 68 classical nights.

Altogether, the dedicated CLAUDS observations resulted in useful images with a total open-shutter time of $t_{\text{exp}} = 280.60 \text{ h}$ (i.e.

not counting overheads and weather or instrument problems); of this, 219.75 h were taken with the new u filter, and 60.85 with the old u^* . Additionally, we reprocessed and incorporated into our stacks 181.48 h of archival u^* data, including 51.62 h of u^* obtained in two ultradeep $\sim 1 \text{ deg}^2$ MegaCam pointings as part of the MUSUBI program. In total, the full stacked images (dedicated CLAUDS plus archival data) represent 462.09 h of open-shutter time, which is equivalent to ~ 112 classical-mode nights (after accounting for overheads, telescope faults, and weather).

Dedicated CLAUDS observations were carried out in CFHT’s queue mode during dark-time runs. The bulk of the data was obtained using 600 s exposures taken in groups of either six or 11 dither positions (64 and 117 min blocks of time, respectively, including overheads). We used the pre-defined MegaCam Large Dithering Patterns LDP6 and LDP11, which cover a $30 \text{ arcsec} \times 180 \text{ arcsec}$ ellipse² and are designed to fill in the interchip gaps as well as to give good data for constructing sky flats and to ensure satisfactory masking of the few bad detector pixels. A very small number of shorter, 180 s exposures was also taken at positions that overlapped more than one of the main fields for the purpose of photometrically tying the fields together, if needed. Where suitable, we also used archival MegaCam u^* data – including those from the MUSUBI program – all retrieved from the CFHT archive at the Canadian Astronomy Data Centre³ (CADC). Two subfields of the XMM-LSS field (see Table 2) are based heavily on such archival data, while in the COSMOS field we used archival and MUSUBI data to construct the u^* stacks. (The u data in E-COSMOS, and in the other fields, were all taken as new observations under the CLAUDS observing program.) The archival data were taken from programs executed between 2003 December 22 and 2016 March 31. These data were taken with a variety of observing strategies but with multipoint dither patterns and typical exposure times between 300 and 660 s (though a small number had shorter exposure times).

3 DATA REDUCTION AND QUALITY

3.1 Data reduction

3.1.1 Image processing, calibration, and matching with the HSC-SSP data

As the first step in data reduction, the individual MegaCam images are pre-processed by the ELIXIR software (Magnier & Cuillandre 2004) at CFHT. ELIXIR applies detrending steps, namely overscan correction, bias subtraction, flat-fielding, and masking, before the data are transferred to CADC for further processing.

Subsequent processing, namely astrometric/photometric calibration and image stacking, is done using the MEGAPIPE data pipeline (Gwyn 2008) at the CADC. The following is a summary of this MEGAPIPE procedure, highlighting some of the modifications that were made to the original pipeline to accommodate the CLAUDS data.

Since the Gaia astrometric catalogue (Brown et al. 2016) became available, MEGAPIPE has used it as an astrometric reference frame. However, early HSC astrometry, including that of Public Data Release 1 (PDR1; Aihara et al. 2018b) and internal data release S16A was tied to a pre-Gaia version of Pan-STARRS (Magnier et al.

¹http://cfht.hawaii.edu/en/science/LargePrograms/LP_18B_22A/ImplementationProcedures

²For dither positions and other details see <https://www.cfht.hawaii.edu/Instruments/Imaging/Megacam/specsinformation.html>

³<http://www.cadc-ccda.hia-ihc.nrc-cnrc.gc.ca>

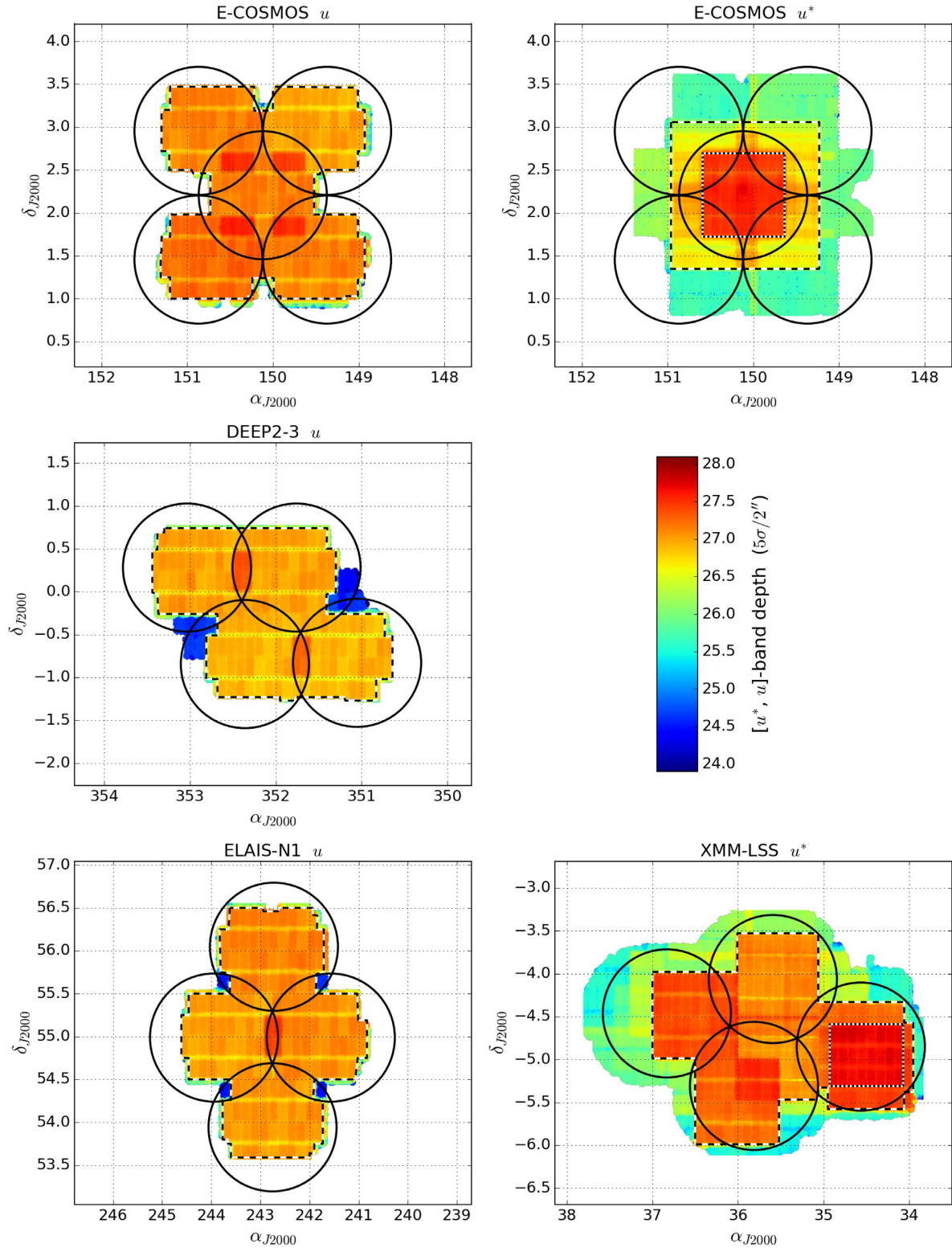


Figure 3. Depth and areal coverage of the CLAUDS data. The E-COSMOS field contains data in two U -band filters (u^* and u) which are stacked separately; these are shown as two panels in this Figure. The black circles represent the nominal Subaru HSC pointings of the HSC-SSP. The dashed lines mark the extent of the CLAUDS Deep data, and dotted lines that of the CLAUDS UltraDeep. See Section 3.2.1 for details of the depth measurements.

Table 2. CLAUDS fields and subfields.

Field and subfield	Other name	RA centre (J2000)	Decl. centre (J2000)	Filter	Notes
XMM-LSS 0	SXDS	02:18:15.60	−04:51:00.0	u^*	Archival+MUSUBI data
XMM-LSS 1	–	02:22:11.00	−05:00:00.0	u^*	–
XMM-LSS 2	–	02:22:11.00	−04:03:00.0	u^*	–
XMM-LSS 3	–	02:25:59.00	−04:29:40.0	u^*	–
XMM-LSS 4	–	02:24:03.00	−04:29:40.0	u^*	–
XMM-LSS 5	CFHTLS-D1	02:25:59.00	−04:30:00.0	u^*	Archival data
E-COSMOS SW	–	09:58:03.00	+01:27:21.0	u	–
E-COSMOS NE	–	10:02:50.20	+02:57:21.0	u	–
E-COSMOS NW	–	09:58:03.00	+02:57:21.0	u	–
E-COSMOS SE	–	10:02:50.20	+01:27:21.0	u	–
E-COSMOS C	CFHTLS-D2 / COSMOS	10:00:28.60	+02:12:21.0	u	–
				u^*	Archival+MUSUBI data
ELAIS-N1 N	–	16:11:00.00	+56:00:00.0	u	–
ELAIS-N1 S	–	16:11:00.00	+54:00:00.0	u	–
ELAIS-N1 W	–	16:07:35.00	+55:00:00.0	u	–
ELAIS-N1 E	–	16:14:25.00	+55:00:00.0	u	–
DEEP2-3 SE	–	23:28:47.00	−00:45:55.0	u	–
DEEP2-3 NW	–	23:27:37.00	+00:13:50.0	u	–
DEEP2-3 SW	–	23:24:53.00	−00:45:55.0	u	–
DEEP2-3 NE	–	23:31:31.00	+00:13:50.0	u	–

2016), which in turn was tied to 2MASS (Skrutskie et al. 2006). In contrast, more recent versions of HSC-SSP data, including PDR2 (Aihara et al. 2019), use *Gaia* astrometry. Small but significant shifts, of the order of 0.1 arcsec, exist between these two reference frames and may be important for some science applications. We therefore produce two versions of the CLAUDS image stacks, one matched to pre-*Gaia* Pan-STARRS astrometry, and one to *Gaia* astrometry. The MegaCam data used in Section 4 of this paper use the pre-*Gaia* version of Pan-STARRS to match the reference frame of the HSC internal data release we used here. The *Gaia*-calibrated CLAUDS stacks should, of course, be used with the more recent, *Gaia*-calibrated HSC images. The internal astrometric accuracy within the MegaCam data is better than 0.04 arcsec RMS, and the HSC and MegaCam images are aligned to each other with the same accuracy, provided that the pre-*Gaia* or post-*Gaia* stacks are used consistently.

The photometric calibration of CLAUDS is tied to the SDSS. The SDSS photometry is transformed into the the MegaCam system using the following transformations:

$$u^* = u_{\text{SDSS}} - 0.241 (u_{\text{SDSS}} - g_{\text{SDSS}})$$

$$u = u_{\text{SDSS}} + 0.036 (u_{\text{SDSS}} - g_{\text{SDSS}}) - 0.165, \quad (1)$$

where u and u^* are the CFHT filters while u_{SDSS} and g_{SDSS} refer to SDSS filters. These transformations are derived using synthetic photometry. To achieve this, the full response function of each filter is computed, including the transmission of the filter itself, the CCD quantum efficiency, the transmission of the MegaPrime optics, the reflectance of the CFHT primary mirror, and 1.25 airmasses of atmospheric attenuation. These response functions are multiplied by stellar spectra from Pickles (1998) and Bohlin, Gordon & Tremblay (2014, the CALSPEC spectra) – the synthetic star colours thus produced are shown as blue and red points in Fig. 4. For comparison, colours of real stars in the COSMOS field are shown as black points. We ignore the effects of Galactic extinction since the differential extinction between the U filters (u^*/u and u_{SDSS}) is very small in this high-latitude fields. The fits given by equation (1) are shown as a green lines. Note that the transformations are only valid for stars with $u_{\text{SDSS}} - g_{\text{SDSS}} > 1.2$.

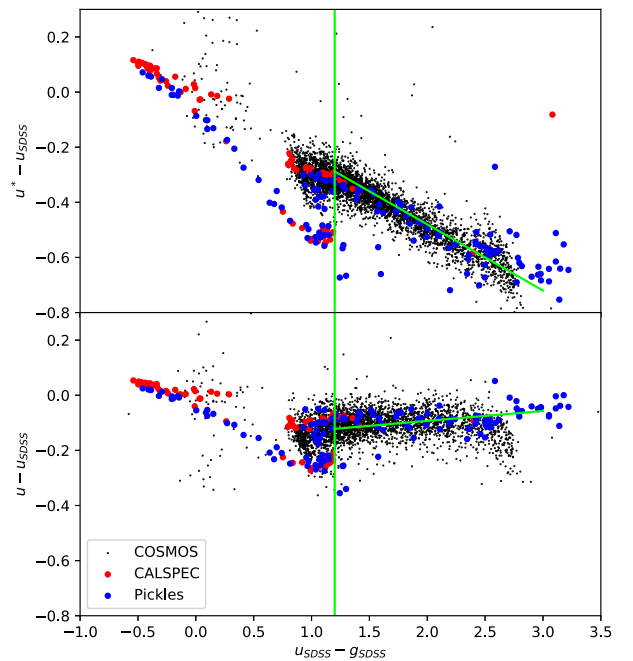


Figure 4. Transformations from the SDSS to the CFHT u -band filters. Photometry from the COSMOS field is shown as black dots. Synthetic photometry based on the CALSPEC (Bohlin et al. 2014) and Pickles (1998) spectra are shown in red and blue, respectively. The green line shows the adopted transformations between the SDSS and the CFHT U filters, u^* and u . The fits only take into account data redwards of $u_{\text{SDSS}} - g_{\text{SDSS}} = 1.2$, as indicated by the vertical line.

When calibrating with the SDSS, we use only sources classified as stars by SDSS (using the SDSS parameter `probPSF`). We use the SDSS PSF magnitudes. The corresponding instrumental MegaCam magnitudes are measured through circular apertures whose size is set by the seeing of each image. Two apertures are used, with diameters 2 and 5.15 times the mean FWHM of stars in the image. The large aperture magnitude is equivalent to a Kron magnitude for

point sources in MegaCam images. However, this aperture is quite large and consequently produces noisy measurements. The brighter stars in each image are used to determine an offset between the two sets of aperture magnitudes. The flux is then measured through the smaller aperture for all sources and corrected to the larger aperture. The zero-points computed in this way are useful for both stellar and galactic photometry.

For each MegaCam observing run, MEGAPIPE builds a map of the differential zero-point offsets across the detector mosaic, using all the u (or u^*)-band images which overlap the SDSS. This includes a large number of images that are not on CLAUDS fields, ensuring that the differential zero-point corrections are not affected by any local errors in the SDSS. The differential corrections are computed on a CCD-by-CCD basis. This is adequate most of the time, with the caveat that for some runs there is some evidence that the four corner MegaCam chips have cross-chip gradients of about 1–2 per cent. The differential corrections change somewhat between observing runs, with occasional larger changes due to changes in the ELIXIR recipe.

Photometricity is determined using CFHT’s SKYPROBE camera which monitors sky transparency once per minute by observing a set of standard stars located in the direction the telescope is pointing (Cuillandre et al. 2002). For each photometric night, MEGAPIPE measures a nightly zero-point, again using all available images which overlap with the SDSS. CLAUDS images taken on photometric nights are used to build a catalogue of photometric standards in the CLAUDS fields. This catalogue is then used to calibrate all the images, including those taken on non-photometric nights.

The photometric calibration of the individual images is cross-checked before stacking. After the stacks are generated, their photometry is checked against that of the individual images. The calibration is found to be self-consistent to 0.005 magnitudes RMS. Both the input images and the stacks are compared back against the SDSS and we find that typical offsets are 0.015 magnitudes. Note that although we use the SDSS as a calibration reference, we do not directly use the individual stars in the CLAUDS fields as standards, so this last test is effectively independent.

The astrometric/photometric calibrations computed above are stored in external header files. Once the input images are calibrated, they are stacked as follows, noting that u and u^* data are stacked separately where they overlap (such as in the E-COSMOS field). To match the HSC pixel scale and image format, we produce stacks that cover exactly the 4200×4100 pixel tiles corresponding to the ‘patches’ in the HSC-SSP data (Aihara et al. 2018b); i.e. full mosaics covering the whole field are not constructed. The tiles measure $4200 \times 4100 \times 0.168 \text{ arcsec} \times 0.168 \text{ arcsec}$ pixels or about 0.2 degrees on the side. To generate a tile, MEGAPIPE determines which CCDs of which images overlap a particular patch. The CADC storage system allows individual FITS image extensions corresponding to the MegaCam CCDs to be extracted from the telescope archives with very little overhead. Therefore, only the CCDs relevant to a given patch are retrieved, greatly speeding the processing. These input images are resampled on to the HSC pixel grid using SWARP (Bertin et al. 2002) and a Lanczos-3 kernel. Background subtraction is done with SWARP using a 128×128 pixel mesh. Finally, the images are combined by using a mean with 3σ clipping. Weight (inverse variance) maps for the images are also produced with the same format and pixel scale. The variance maps are produced by summing the variance maps of the individual input images, which in turn are derived from the sky in those images.

The end product is a set of u and/or u^* image tiles that have the same dimensions and pixel scale, and are astrometrically aligned with, HSC-SSP *grizy* image tiles. The images also have corresponding weight (inverse variance) maps. As is the case for the HSC images, there is a small overlap between tiles which is useful in avoiding boundary effects in object detection and photometry.

To give a visual impression of the quality of both the CLAUDS and HSC-SSP data, in Fig. 5 we show a 1 arcmin \times 1 arcmin subregion within the central COSMOS field. Here, the u depth is typical of the entire CLAUDS data set while the HSC *grizy* images are at the depth that the whole HSC-SSP Deep Layer will reach upon completion of Subaru observations, expected in 2021.

3.1.2 Photometry and merged catalogues

The steps described in 3.1.1 give us u/u^* images that are aligned with the HSC-SSP *grizy* images. The next step is to produce combined multiband catalogues. For this, we use two completely independent, parallel procedures and produce two sets of independent catalogues.

The first procedure applies the SEXTRACTOR software (Bertin & Arnouts 1996) to a merged *Ugrizy* image constructed from combining the CFHT u/u^* and Subaru *grizy* images. The second procedure employs a modified version of the HSC data reduction pipeline (HSCPIPE; Bosch et al. 2018). Both procedures will be described in detail in forthcoming papers (A. Golob et al. in preparation; J. Coupon et al. in preparation), so here we only briefly describe their key elements. Having two independent photometric pipelines will let us in a future paper validate the relatively new HSCPIPE applied to our U -band data against the simpler but well-established approach that SEXTRACTOR represents.

In both cases presented here we use our CLAUDS data as described in Section 3.1.1 and images from the S16A internal release of the HSC-SSP data that are deeper than those from the first HSC-SSP Public Data Release (PDR1; Aihara et al. 2018b). As the HSC data accumulate, future updates of our catalogues will incorporate newer HSC-SSP data releases.

i. SEXTRACTOR-based catalogues: We produce two catalogues using SEXTRACTOR: (1) a single-band U catalogue (keeping u and u^* separate) and (2) a multiband catalogue with object detection performed on a combined *Ugrizy* image.

The U -band catalogue is created running SEXTRACTOR on the u/u^* images in the simple single-image mode with detection parameters tuned for this data set. We record various measurements for each detected source, including position, fluxes (in Kron, isometric, and fixed-radius circular apertures), fiducial radius, ellipticity, position angle, and central surface brightness.

For the multiband catalogue, combined *Ugrizy* images, ΣSNR , are created for each of the HSC patches by combining the observations in the N bands observed as

$$\Sigma\text{SNR} = \sum_{i=1}^N \left(\frac{f_i - \mu_i}{\sigma_i} \right), \quad (2)$$

where f_i is the flux in each pixel, σ_i is the RMS width of the background sky distribution, and μ_i is its mean. The index i runs over MegaCam bands u or u^* (or both, where available – i.e. in the central area of E-COSMOS) as well as the HSC bands *grizy*. Like the χ^2 image (Szalay, Connolly & Szokoly 1999) that’s often used for object detection, the ΣSNR image combines information from all available bands. However, in contrast to χ^2 , ΣSNR places more

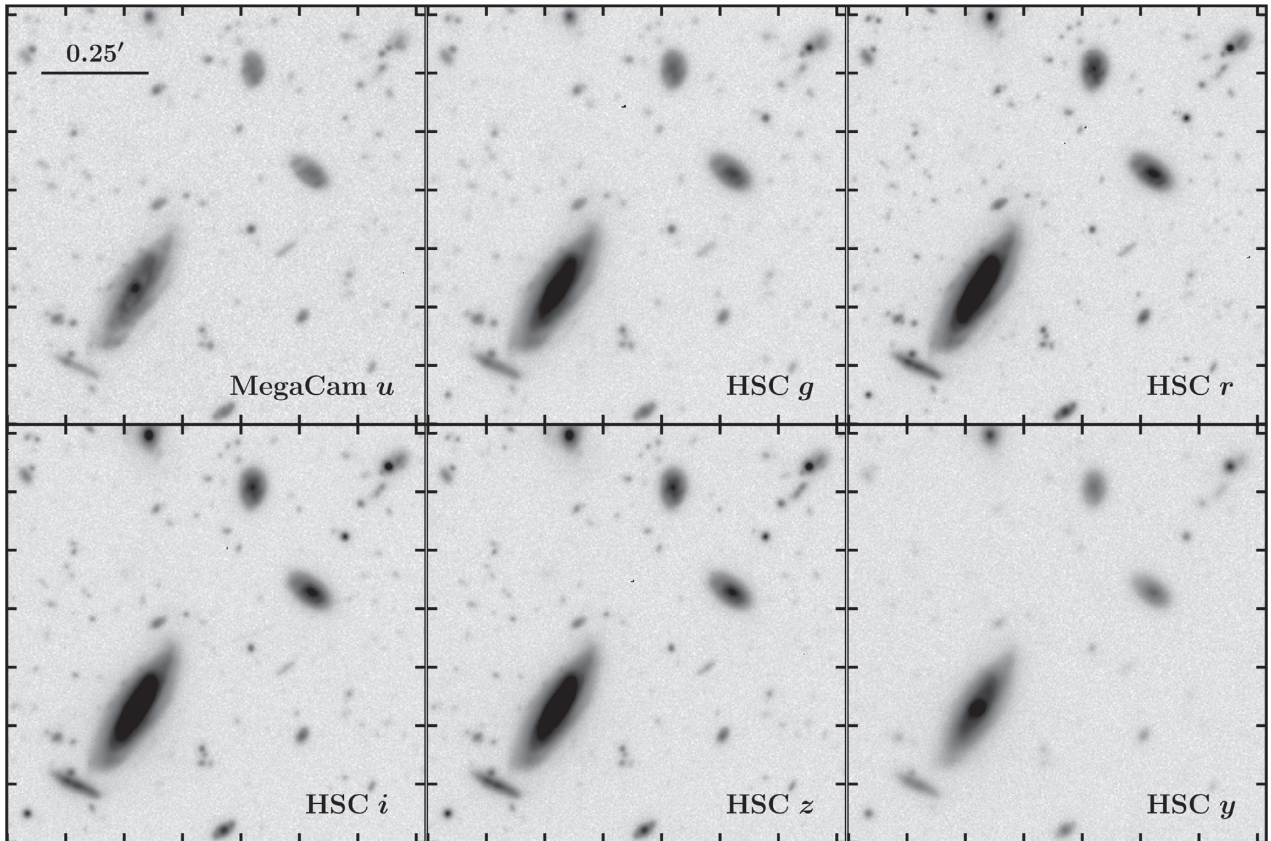


Figure 5. An example of the CLAUDS u and HSC-SSP $grizy$ data at Deep depth. Each panel shows the same representative $1\text{ arcmin} \times 1\text{ arcmin}$ subfield within the central COSMOS field but imaged through a different filter. The u data in this field are of the depth that is typical of most of the CLAUDS survey ($u \approx 27\text{AB}$, 5σ in 2 arcsec apertures), while the $grizy$ data shown here have the depth that will be typical of the HSC-SSP Deep Layer once the HSC-SSP is fully completed. When the HSC $grizy$ observations are completed, the combined surveys will contain 72 000 times the area shown in this $1\text{ arcmin} \times 1\text{ arcmin}$ example, to comparable or deeper depths in $Ugrizy$.

weight on the shallower bands, which can improve the sensitivity to objects with strong colours; is less sensitive than χ^2 to seeing variations between bands (see Bertin et al. 2002); and avoids the possibility of spurious detections due to regions of correlated negative noise that can become positive when squared in the χ^2 image.

The multiband catalogue is next created by running SEXTRACTOR in dual image mode using the combined $Ugrizy$ ΣSNR images (equation 2) for object detection. For each detected source we record various measurements, including position, fluxes in all available bands (in Kron, isometric, and fixed-radius circular apertures), fiducial radii, ellipticities, position angles, and central surface brightnesses. We also record ‘fluxes’ in the ‘ ΣSNR ’ band defined via equation (2); these do not have a physical meaning, but are useful for understanding the detection properties of the catalogue.

The HSC-SSP consists of overlapping rectangular tracts which are divided into square patches. Adjacent HSC-SSP patches, as well as the CLAUDS data which are registered to them (Section 3.1.1) overlap by 200 pixels so we discard all objects whose centres fall within 100 pixels of a patch edge, giving a catalogue of unique sources within a single tract. We merge the resulting tract catalogues by matching objects by position and, where duplicate objects are detected, keeping only the one with the highest S/N in the detection band (u^* , u , or combined $Ugrizy$, depending on the catalogue in question).

ii. HSCPIPE-based catalogues For the second procedure we use a modified version of the HSC pipeline (HSCPIPE; Bosch et al. 2018) to combine the CLAUDS U -band data with the HSC $grizy$ data. We modified HSCPIPE, which is designed to work with HSC images, to allow it to handle our stacked MegaCam u/u^* images. All other features and functionality of HSCPIPE remain unchanged.

To adapt MegaCam data for ingestion into HSCPIPE, we create a CLAUDS ‘exposure’ object (in the form of a multiextension FITS file) that contains the image, the variance map and the mask plane, all registered on to the same pixel grid and sky tiles (i.e. identical tracts and patches) as the HSC images. The variance map is simply the inverse weight map created by MEGAPIPE (Section 3.1.1) and the mask plane is identical to that of the g -band image, which contains the bright-star masks (Coupon et al. 2018) and the geometry of the HSC-SSP survey.

To detect sources and accurately measure their photometry and morphology, the HSC pipeline accounts for point spread function (PSF) variations between the bands and across the image. The full software normally starts from single exposures, fits a spatially varying PSF model from high-SNR point sources, and co-adds the PSF models in the same way as the exposures (for example, using a weighted-mean estimate). Here, for simplicity, we only measure the PSF on the co-added image. It is likely that the discontinuities of the PSF will not be properly reproduced in our best-fitting PSF model, however, given the large number of co-added exposures in

CLAUDS, we do not expect that this will have a significant impact on the measured photometry.

Next, we run the source detection on the u - and u^* -band images only (this would be done during the co-addition process in the nominal version of the HSC pipeline), so that we can finally run the data set process on the full $Ugrizy$ data set. This last step is done using the version of HSCPIPE described by Bosch et al. (2018). In brief, HSCPIPE first merges the footprints of the detections made on individual images, then picks the best-suited image to proceed with the de-blending of sources it considers to be blended (the images are ordered per priority, the highest priority being the i -band image followed by other bands if the i -band SNR is too low), and finally performs a series of measurements (with uncertainty estimates) of the source photometry and morphology. Of particular note is that HSCPIPE's primary photometric output is CModel magnitudes for each band, including CLAUDS u and/or u^* , which are measured by finding the best PSF-convolved morphological model for each object (see Bosch et al. 2018 for details).

3.2 Data quality

3.2.1 Depth

Fig. 3 shows the limiting magnitude (i.e. the depth) of our u - and u^* -band data across each CLAUDS field expressed as 5σ (i.e. $\text{SNR} = 5$) in 2 arcsec diameter apertures. This measurement was derived from the U -selected SEXTRACTOR-based catalogue (Section 3.1.2) on a fine grid of positions by taking the median of the magnitudes of nearby sources that have $\text{SNR} = 5$ in 2 arcsec diameter apertures. This method has the advantage of providing empirical depth measurements with high angular resolution but requires reliable measurements of the photometric uncertainties. The photometric uncertainties were calibrated empirically (see below) and this gave a multiplicative rescaling of the RMS by a factor of 1.5 to produce the final measurements shown in Fig. 3.

The depth measurements shown in Fig. 3 were calibrated using two different estimates of the limiting magnitude, m^{lim} . In the first of these two approaches, we first created two images, each from an independent half of the available exposures (split by the MegaCam exposure number into 'even' and 'odd' stacks with essential an equal number of individual exposures contributing to each stack). The m^{lim} was then derived from the difference of magnitude, Δm , measured for the same objects in the two half-exposure images. Here, the depth of the half-depth stacks, $m_{1/2}^{\text{lim}}$, is obtained by comparing the RMS scatter in the magnitudes of objects detected at $\text{SNR} = 5$. The full depth is then extrapolated through $m^{\text{lim}} = m_{1/2}^{\text{lim}} + 2.5 \log(\sqrt{2})$, where the $\sqrt{2}$ accounts for the difference in exposure time between the full and the half-exposure images. The second estimate of m^{lim} was computed from the scatter of the background noise. To do so, the standard deviation of the background flux σ_{bf} was measured in 2 arcsec apertures randomly placed across the image, while avoiding detected objects. The limiting magnitude is then given by $m^{\text{lim}} = -2.5 \log(\text{SNR} \times \sigma_{\text{bf}}) + ZP$, where $\text{SNR} = 5$ and ZP is the zero-point of the image. The two estimates of m^{lim} agree with one another, after rescaling the raw U -band uncertainties by a factor of 1.5. Such uncertainty underestimates in the raw SNR measurements were expected, especially given the correlation between output pixels that is introduced by our stacking processes (see Section 3.1.1).

In Fig. 3 the dashed lines outline an 18.60 deg² CLAUDS Deep area outside of which the U -band depth falls off rapidly (see also Fig. 6). Within this Deep area, our data reach $U > 26.60$ AB

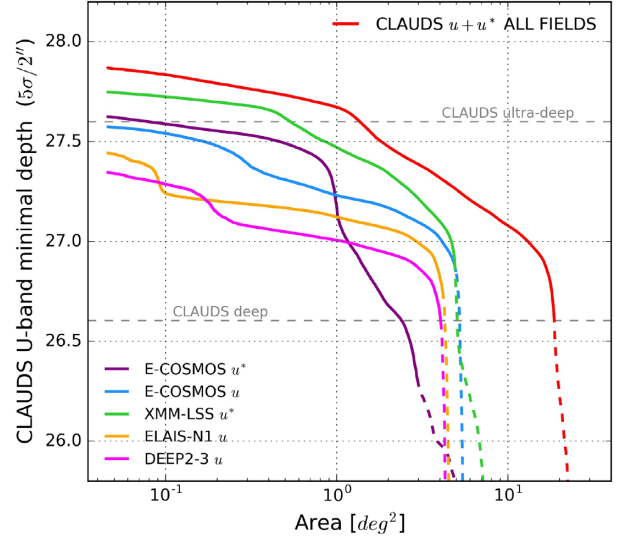


Figure 6. The limiting depth as a function of area for the total CLAUDS (red curve) and the individual constituent fields. The u^* and u are held separate for the E-COSMOS field. The solid curves show the depth inside the footprints outlined in Fig. 3, while the dashed curves are outside those footprints.

(5σ , 2 arcsec apertures) at every position, and the median depth is $U = 27.09$ AB. The CLAUDS median depth is shown in Fig. 1, and the minimal depth as a function of area is shown in Fig. 6. Within the CLAUDS UltraDeep area, defined by the dotted lines in Fig. 3, our data reach $U > 27.60$ AB (5σ , 2 arcsec apertures) at every position, and the median depth over this area is $U = 27.71$ AB.

To our knowledge, these are the deepest images ever taken in the U -band over such large areas.

3.2.2 Seeing

Fig. 7 shows U -band seeing maps measured from the stacked images. To build these maps, we first extract from our U -selected SEXTRACTOR catalogues a sample of high-probability stars with $U < 24$ AB, $\text{S/N} > 5$, and $\text{SEXTRACTOR CLASS_STAR} > 0.9$. At every position on the sky, we use the median of the half-light radii of all objects from this sample within a distance of 0.05 deg to characterize the local PSF. Finally, we convert the half-light radii to the PSF FWHM using a conversion factor derived empirically from measurements of the MegaCam PSF (Gwyn 2008).

The median seeing over the entire deep area of CLAUDS is 0.92 arcsec but varies across the survey, as can be seen in Fig. 7. Nevertheless, it is typically $\lesssim 1$ arcsec over the survey and is substantially better in some subfields.

4 ILLUSTRATIVE APPLICATIONS

In this section we show three initial illustrative applications of our data set: galaxy U -band number counts, U -band dropout number counts, and photometric redshifts. All three results are preliminary and will be refined in separate papers (A. Golob et al., in preparation; C. Liu et al., in preparation).

4.1 U -band galaxy number counts

Galaxy number counts are a standard way to characterize the observed galaxy population and also to test a new data set and

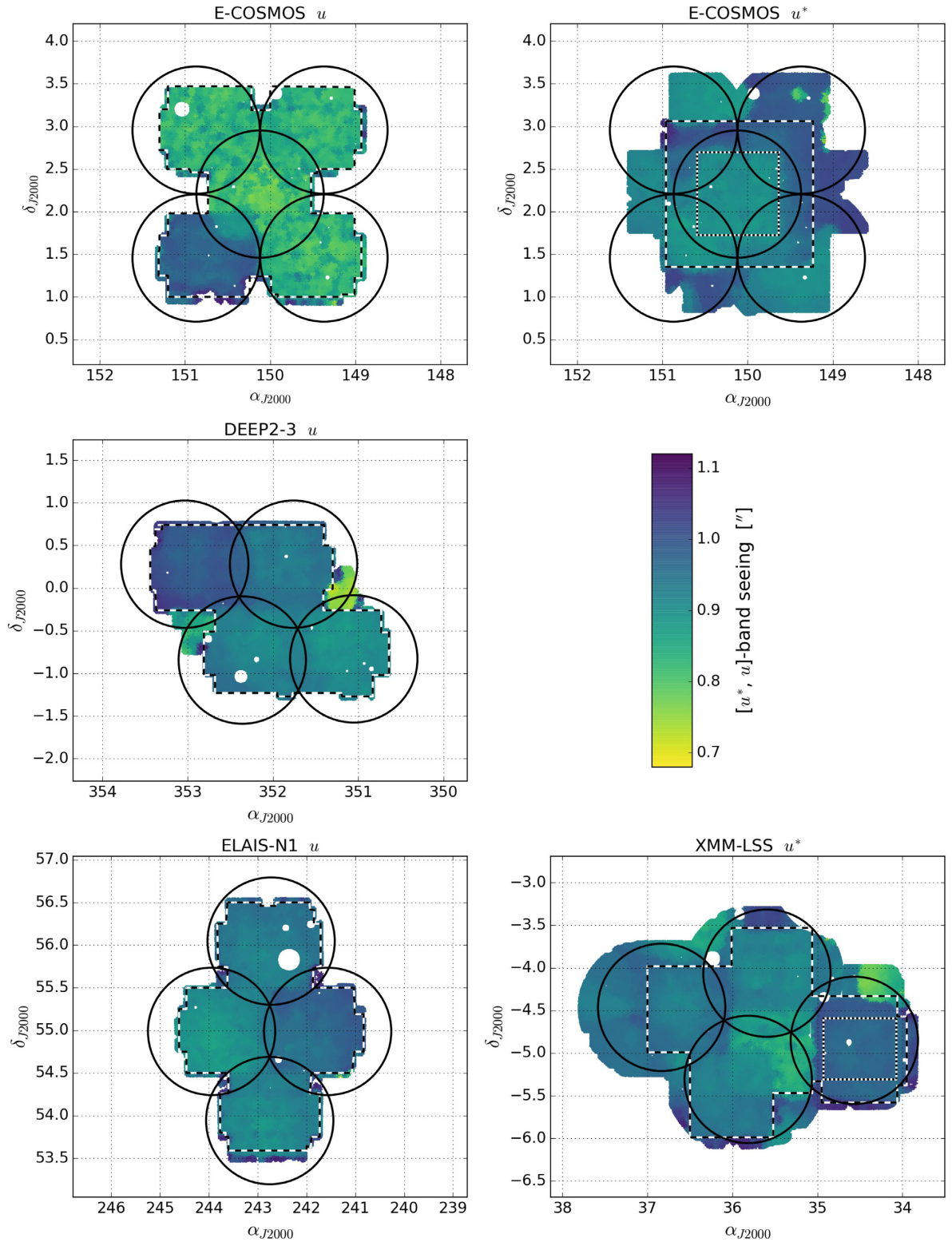


Figure 7. Seeing in the stacked images measured from stellar half-light radii across the CLAUDS fields. The black circles represent the nominal Subaru HSC pointings of the HSC-SSP. The dashed lines mark the extent of the CLAUDS Deep data, and the dotted lines that of the CLAUDS UltraDeep. See Section 3.2.2 for details of the seeing measurements.

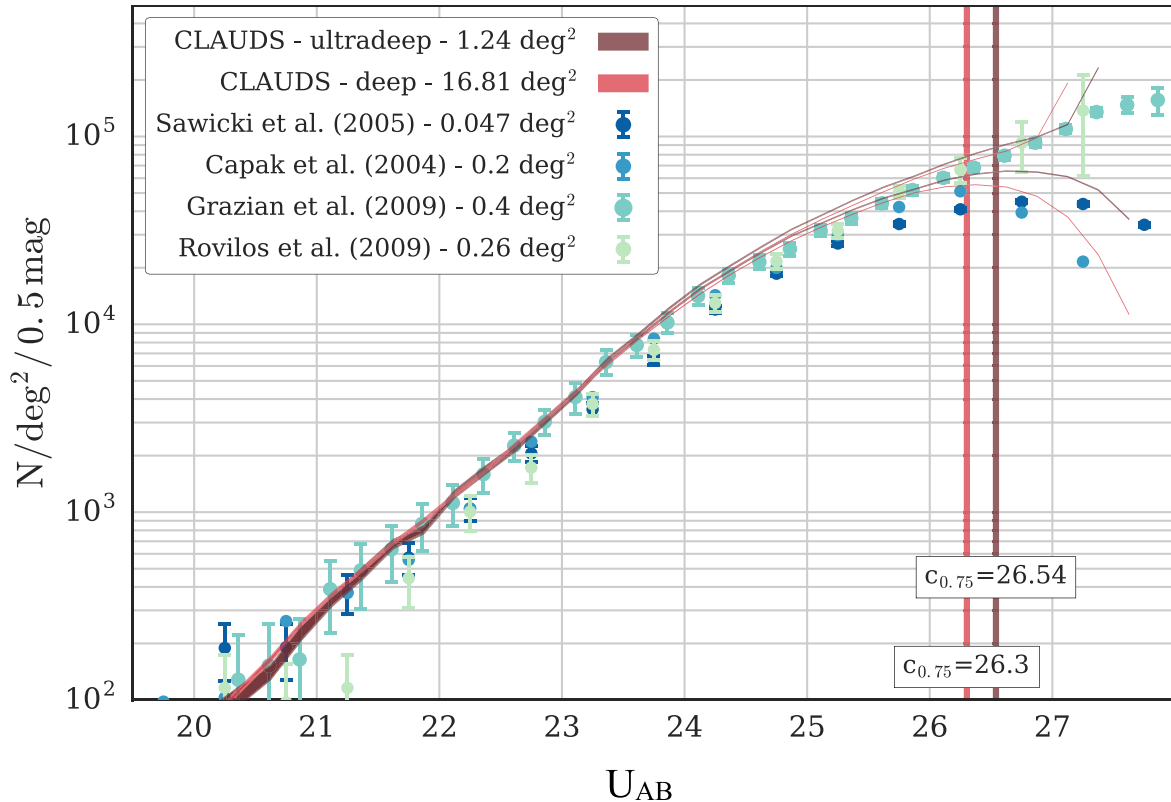


Figure 8. CLAUDS U -band galaxy number counts compared with other deep surveys. The CLAUDS counts are from SEXTRACTOR object detection and photometry performed on U -band images only, as described in Section 3.1.2. We found that there was no appreciable difference in the number counts in the u and u^* bands and so we simply combine the data with an inverse-variance weighting. The CLAUDS counts are shown both in the Deep and UltraDeep, after excluding regions masked by bright stars or artefacts. The deep area listed does not include the UltraDeep area. Both raw counts and incompleteness-corrected counts are shown: incompleteness-corrected counts are the higher two curves. The width of the curves corresponds to 68 per cent confidence regions and the vertical lines mark the 75 per cent completeness depths in the raw Deep and UltraDeep counts. U -band counts from Grazian et al. (2009) and Rovilos et al. (2009) are shown after correction for incompleteness by these authors, while the counts from Capak et al. (2004) and Sawicki & Thompson (2005) are raw, uncorrected counts.

gauge its depth. The unprecedented combination of depth and area of CLAUDS allows a measurement of the U -band counts that has excellent statistics over a wide range in magnitude (see Fig. 8) and is largely insensitive to the effects of large-scale structures (i.e. ‘cosmic variance’).

Our CLAUDS number counts are derived using the single band (u - or u^* -selected) SEXTRACTOR catalogue. Effective areas are calculated by randomly sampling the image masks at one million points per deg^2 . Stars were removed from the sample used to compute the number counts using SEXTRACTOR’s CLASS_STAR parameter with a threshold of 0.85 and measured fluxes were corrected for Galactic extinction using the prescription of Schlegel, Finkbeiner & Davis (1998). U -band number counts and confidence intervals were computed for each area with 1000 bootstrap resamplings of the catalogue before a histogram was created. In each bootstrap iteration, the flux of each selected object was drawn from its Gaussian probability density function (PDF).

The completeness of the catalogue was determined by adding synthetic objects to each of the detection images and measuring the recovery rate of objects as a function of flux. These artificial objects were generated using SEXTRACTOR’s object-modelling functionality during the creation of the U -selected catalogues. As such, each artificial object has a real counterpart in the U -band catalogues and has model properties (shape, size) that are representative of the real object. The additional crowding introduced by doubling

the number of objects in the detection images could impact the completeness measurements; to minimize this bias, we record the distance of each real object in the original catalogue to its nearest neighbour and only use those SEXTRACTOR model objects whose nearest neighbour in the combined (model+original) catalogue is no closer than the nearest neighbour of its original counterpart in the completeness calculations.

Table 3 lists, for each magnitude bin, the median value of the 1000 bootstrap resamplings, along with the 16–85 per cent completeness intervals; the values are shown graphically in Fig. 8 which shows the number counts calculated over the CLAUDS Deep and UltraDeep areas (red and brown shaded regions, respectively) and their incompleteness-corrected counterparts. Stellar objects have been removed and magnitudes for the CLAUDS number counts are SEXTRACTOR’s MAG_AUTO measurements. Note that the CLAUDS Deep and UltraDeep areas given in Fig. 8 are after correcting for regions masked due to bright stars and artefacts.

We compare our U -band counts to previous deep U -band surveys (all taken with 8 m class telescopes: Capak et al. 2004; Sawicki & Thompson 2005; Grazian et al. 2009; Rovilos et al. 2009) and find our result to be in good agreement with these earlier measurements. CLAUDS 75 per cent U -band detection completeness, determined from our simulations, is shown with vertical lines in the figure and is $c_{0.75} = 26.30$ AB in the CLAUDS Deep area and $c_{0.75} = 26.56$ AB in the UltraDeep.

Table 3. CLAUDS raw U -band galaxy number counts for the Deep and UltraDeep regions of the survey. Counts are raw counts (i.e. not corrected for incompleteness) and are given as the logarithm of $N \text{ deg}^{-2} 0.5 \text{ mag}^{-1}$. The 16–84 percentile uncertainty range is also given. The estimated completeness (columns 5 and 9) can be used to convert the raw counts to incompleteness-corrected counts.

U (AB)	Counts $\log(N/\text{deg}^2/0.5 \text{ mag})$ (Deep)	16 percent (Deep)	84 percent (Deep)	Completeness (Deep)	Counts $\log(N/\text{deg}^2/0.5 \text{ mag})$ (UD)	16 percent (UD)	84 percent (UD)	Completeness (UD)
19.125	0.3059	0.1893	0.3976	1.0140	–	–	–	–
19.375	0.8607	0.7997	0.9143	1.0118	0.8110	0.5100	0.9871	1.0144
19.625	1.2902	1.2544	1.3234	1.0093	1.1632	0.9871	1.3229	1.0144
19.875	1.5939	1.5696	1.6157	1.0065	1.6861	1.6041	1.7550	1.0141
20.125	1.8816	1.8629	1.8989	1.0019	1.8717	1.8000	1.9571	1.0116
20.375	2.0518	2.0382	2.0657	0.9994	2.0285	1.9724	2.0732	1.0070
20.625	2.2137	2.2021	2.2265	0.9991	2.1434	2.1066	2.1821	1.0042
20.875	2.4017	2.3928	2.4111	0.9988	2.3535	2.3156	2.3965	1.0014
21.125	2.5574	2.5498	2.5651	0.9965	2.5228	2.4964	2.5517	0.9968
21.375	2.6902	2.6834	2.6971	0.9940	2.6751	2.6400	2.7005	0.9943
21.625	2.8318	2.8259	2.8373	0.9914	2.8259	2.8033	2.8476	0.9937
21.875	2.9511	2.9459	2.9557	0.9892	2.9031	2.8838	2.9195	0.9912
22.125	3.0863	3.0822	3.0908	0.9886	3.0909	3.0739	3.1067	0.9863
22.375	3.2124	3.2089	3.2162	0.9863	3.2171	3.2005	3.2327	0.9815
22.625	3.3566	3.3536	3.3596	0.9838	3.3362	3.3236	3.3479	0.9790
22.875	3.4976	3.4951	3.5003	0.9813	3.4837	3.4733	3.4968	0.9787
23.125	3.6400	3.6378	3.6424	0.9787	3.6318	3.6226	3.6399	0.9784
23.375	3.7855	3.7835	3.7875	0.9762	3.8003	3.7949	3.8065	0.9759
23.625	3.9192	3.9176	3.9206	0.9736	3.9276	3.9223	3.9346	0.9711
23.875	4.0438	4.0423	4.0452	0.9708	4.0644	4.0588	4.0695	0.9660
24.125	4.1607	4.1594	4.1620	0.9657	4.1863	4.1820	4.1908	0.9609
24.375	4.2655	4.2645	4.2667	0.9580	4.2865	4.2827	4.2905	0.9555
24.625	4.3594	4.3585	4.3604	0.9481	4.3818	4.3786	4.3856	0.9476
24.875	4.4458	4.4449	4.4467	0.9376	4.4706	4.4683	4.4737	0.9351
25.125	4.5203	4.5194	4.5211	0.9249	4.5401	4.5380	4.5430	0.9201
25.375	4.5872	4.5865	4.5879	0.9093	4.6108	4.6084	4.6128	0.9045
25.625	4.6492	4.6485	4.6499	0.8881	4.6720	4.6694	4.6745	0.8864
25.875	4.7004	4.6998	4.7010	0.8544	4.7188	4.7168	4.7213	0.8632
26.125	4.7322	4.7316	4.7328	0.8040	4.7680	4.7661	4.7704	0.8318
26.375	4.7432	4.7426	4.7438	0.7389	4.8000	4.7979	4.8019	0.7865
26.625	4.7311	4.7305	4.7317	0.6451	4.8160	4.8141	4.8179	0.7298
26.875	4.6816	4.6809	4.6822	0.4795	4.8106	4.8085	4.8125	0.6543
27.125	–	–	–	–	4.7861	4.7842	4.7881	0.5156

4.2 Photometric redshifts

Given the depth and large area of CLAUDS and HSC-SSP imaging, spectroscopy will not be possible for the vast majority of objects even in the eras of the prime focus spectrograph (PFS; Takada et al. 2014) now under construction for Subaru and the planned Mauna Kea spectroscopic explorer (MSE; McConnachie et al. 2016). Quality photometric redshifts are therefore essential for many science applications. Photometric redshifts rely on the detection of prominent spectral breaks in multiband photometric measurements, and here U -band fluxes are essential for localizing the Balmer and 4000 Å breaks at lower redshifts, and the Lyman break at higher redshifts (e.g. Connolly et al. 1995; Gwyn & Hartwick 1996; Sawicki et al. 1997). In this section we illustrate how the combination of CLAUDS U and HSC-SSP $grizy$ observations can give excellent photometric redshifts.

Our team is using several photometric redshift techniques and these will be described in future papers. For the illustration in this paper, our photometric redshifts are computed using a colour-space nearest-neighbour machine learning technique that will be described in a forthcoming paper (A. Golob et al., in preparation). Briefly, we use the 30 band COSMOS photometric redshifts from Laigle et al. (2016) as a training set. We match our SEXTRACTOR-based uu^*grizy catalogue (Section 3.1.2) to that of Laigle et al.

(2016) by object positions with a tolerance of 1 arcsec. For each object in our full catalogue, we calculate colours from MAG_{ISO} photometry and identify the 50 nearest neighbours in colour space with matched COSMOS photo- z 's. We use these 50 nearest neighbours to fit a weighted Gaussian kernel density estimator (KDE), with each neighbour's redshift weighted by $(d_{NN} \times \Delta z)^{-1}$, where d_{NN} is the Euclidean distance in feature (i.e. colour) space to the object under consideration, and Δz is the width of the 68 percent confidence interval of the neighbour's redshift in the Laigle et al. (2016) catalogue. Where $d_{NN} = 0$, we set weights to 0 to prevent objects in the COSMOS field from using *themselves* in their KDEs.

Adding the i -band magnitude to the feature space used to select each object's nearest neighbours effectively acts as a prior probability on the object's redshift: when it is included, objects with brightnesses that differ significantly from the object under consideration are moved to larger distance and penalized in the KDE if included. This is helpful in minimizing outliers at $z_{spec} \sim 3$, but worsens the performance of the estimator at low redshift.

Tanaka et al. (2018) presented photometric redshifts computed from HSC-SSP $grizy$ (but no U) data using several techniques. However, we do not compare our photometric redshifts with their

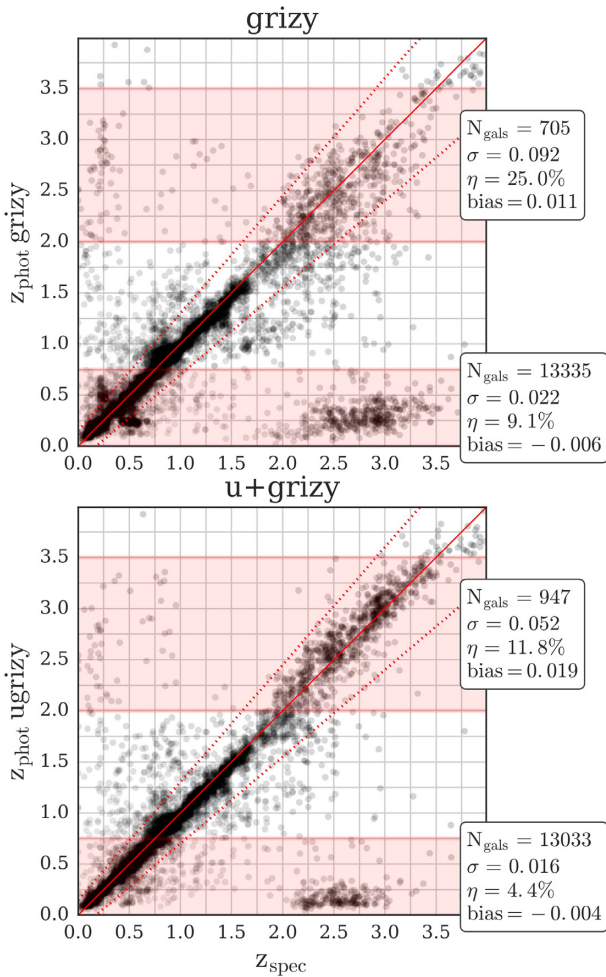


Figure 9. Photometric redshifts without (top) and with (bottom) U -band. The addition of U -band (in this case u^*) to the *grizy* data set significantly improves photometric redshift performance. Scatter (σ , defined as the normalized median absolute deviation, NMAD), outlier fraction (η) and bias are all strongly reduced, particularly at $z < 0.75$ and $2 < z < 3.5$. The statistics presented in the figure are for the two highlighted bands in z_{phot} . The bias at $z \sim 2.5$ in the bottom panel is due to a similar bias that our empirical method inherited from the Laigle et al. (2016) photometric redshifts that we use as our training set; additionally, many of the catastrophic outliers at $z_{\text{spec}} \sim 2.5$ in the bottom panel are due to insufficient training and with further refinements of the method (in preparation) we can eliminate ~ 75 per cent of them.

results as our goal here is to test how the addition of U -band improves photometric redshifts. To do so, we should use the same photometric redshifts code on catalogues that do or do not include the U band. We thus compare photometric redshift results produced by our machine learning procedure applied to *grizy* data alone and to *grizy* combined with U -band. We do the comparison in the E-COSMOS C field where we have both u^* and u data (which we treat as separate measurements), although the two filters are sufficiently similar to each that we are justified in treating this test as one that tests photo- z improvement brought about by the addition of any deep U -band to the HSC-SSP broad-band data.

As can be seen in Fig. 9, the inclusion of the U -band data dramatically improves the quality of photometric redshifts. In this figure we compare our photometric redshifts calculated without U -band (top panel) and with (bottom panel) against 22 005 high-

quality spectroscopic redshifts we extracted from a compilation of surveys in our fields (Lilly et al. 2007; Bradshaw et al. 2013; Le Fèvre et al. 2013; McLure et al. 2013; Comparat et al. 2015; Kriek et al. 2015; Silverman et al. 2015; Masters et al. 2017; Scodreggio et al. 2018; Tasca et al. 2017, and D. Masters, private communication). Of these redshifts, 13 779 have $z_{\text{spec}} < 0.75$ (604 of these are stars and were excluded from the sample) and 1213 are at $2 < z_{\text{spec}} < 3.5$. It is clear that the addition of U -band improves photometric redshift performance, particularly at $z_{\text{phot}} < 0.75$ and $2 < z_{\text{phot}} < 3.5$ (highlighted bands in the figure).

To quantify the improvement due to the addition of U -band data, we calculate the scatter, σ , the catastrophic outlier fraction, η , and the bias in our photometric redshifts. We define σ as the normalized median absolute deviation (NMAD): $\sigma = 1.48 \times \text{median}(\Delta z / (1 + z_{\text{spec}}))$, where Δz is the absolute difference between z_{spec} and z_{phot} . The outlier fraction η is defined as the fraction of galaxies with $\Delta z > 0.15 \times (1 + z_{\text{spec}})$. Finally, the bias is defined as the median value of $(z_{\text{phot}} - z_{\text{spec}}) / (1 + z_{\text{spec}})$.

The addition of U to the *grizy* gives improvements in photo- z performance over the full redshift range. Over the full redshift range, the addition of U -band reduces scatter from $\sigma_{\text{grizy}} = 0.023$ to $\sigma_{U\text{grizy}} = 0.020$, the bias from $-0.006 \times (1 + z)$ to $-0.004 \times (1 + z)$, and the outlier fraction from $\eta_{\text{grizy}} = 5.3$ per cent to $\eta_{U\text{grizy}} = 5.2$ per cent. The improvements are particularly impressive at $z_{\text{phot}} < 0.75$ and $2 < z_{\text{phot}} < 3.5$. At $z_{\text{phot}} < 0.75$, adding the U band reduces the measured scatter from $\sigma_{\text{grizy}} = 0.022$ to $\sigma_{U\text{grizy}} = 0.016$ and the outlier fraction from $\eta_{\text{grizy}} = 9.1$ per cent to $\eta_{U\text{grizy}} = 4.4$ per cent with negligible change to the bias. At $2 < z_{\text{phot}} < 3.5$ the scatter drops from $\sigma_{\text{grizy}} = 0.092$ to $\sigma_{U\text{grizy}} = 0.052$ and the outlier fraction drops from $\eta_{\text{grizy}} = 25$ per cent to $\eta_{U\text{grizy}} = 11.8$ per cent while the bias increases from 0.011 to 0.019, reflecting systematic errors inherited from the Laigle et al. (2016) photometric redshifts that we used as our training set. These improvements are fully consistent with the fact that photometric redshifts rely on the straddling of spectral breaks with photometric bands: as is well known, U fluxes are needed to straddle the Balmer and 4000 Å breaks at lower redshifts, and the Lyman break at higher redshifts (e.g. Connolly et al. 1995; Gwyn & Hartwick 1996; Sawicki et al. 1997).

We note that our empirical photo- z technique is trained using the Laigle et al. (2016) photometric redshifts which in turn use the LEPHARE photo- z code (Arnouts et al. 2002) calibrated using spectroscopic redshifts. Many of the same spectroscopic redshifts also feature in the comparisons shown in Fig. 9 and this might lead one to suspect there is a circularity involved in our photo- z performance statistical measures. However, the LePhare photo- z calibration is very simple: it is only used to adjust their photometry zeropoints globally. Therefore, this means that there are no biases in our photo- z statistics introduced by this approach; moreover, the improvement in *Ugrizy* over *grizy* photometric redshift sets is real as both sets are calculated with the same procedure and spectroscopic samples.

4.3 U -band Drop-outs at $z \sim 3$

The Lyman break technique (Guhathakurta, Tyson & Majewski 1990; Steidel et al. 1996) gives a well-established way to identify and study large samples of star-forming Lyman break galaxies (LBGs) at high redshift (e.g. Steidel et al. 1998, 1999; Sawicki & Thompson 2006b; Hildebrandt et al. 2009a; Ono et al. 2018, and many others). Selection of $z \sim 3$ LBGs, or U -band drop-outs, requires very deep U -band photometry as well as deep data at two

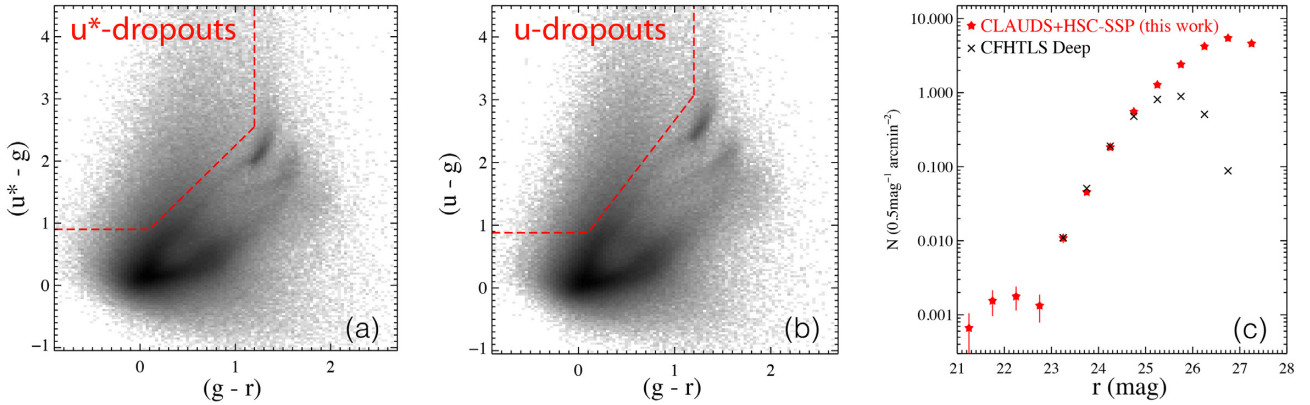


Figure 10. Colour–colour selection of u^* -band dropouts (left-hand panel) and u -band dropouts (middle panel) overlaid on greyscale density plots of objects in the CLAUDS+HSC-SSP. Right-hand panel: number counts of CLAUDS U -band dropouts (combined u^* - and u -dropouts). Our data already allow us to go considerably deeper in U -band dropout studies and over a considerably larger area than was possible in the CFHTLS Deep (black crosses: Hildebrandt et al. 2009a).

longer wavelengths (Steidel et al. 1996), and the combination of CLAUDS and HSC-SSP are excellent for this purpose. Here, the large area of CLAUDS+HSC-SSP will yield very large samples of $z \sim 3$ LBGs that will allow not just excellent statistics and insensitivity to cosmic variance, but will also probe a rich variety of environments (e.g. Toshikawa et al. 2016). In this section we illustrate the ability of CLAUDS U data, in conjunction with gr photometry from HSC-SSP, to select $z \sim 3$ U -band dropout LBGs.

We select Lyman break galaxies at $z \sim 3$ using our HSCPIPE-produced catalogues (see Section 3.1.2). We use HSCPIPE catalogues here for commonality with HSC-SSP studies of LBGs at higher redshifts (e.g. Ono et al. 2018). Because CLAUDS consists of u^* and u observations in different fields, we used both these filters for our dropout selection by employing two similar but different selection criteria for u^* -dropouts and u -dropouts.

For the u^* -dropout selection, we use selection criteria similar to those in Hildebrandt et al. (2009a) who used the very same MegaCam u^* filter, and MegaCam g' and r' filters that are similar to the HSC g and r . The colour selection in the u^*gr diagram is described as

$$\begin{aligned} u^* - g &> 0.9, \\ g - r &< 1.2, \\ u^* - g &> 1.5(g - r) + 0.75, \end{aligned} \quad (3)$$

and is shown with the red lines in the left-hand panel of Fig. 10.

As can be seen in Fig. 2, the central wavelength of the new u -band filter is bluer than that of the old u^* filter. To account for this difference, we modified the u^* -dropout selection criteria of equation (3) and define the selection window in the ugr diagram as

$$\begin{aligned} u - g &> 0.88, \\ g - r &< 1.2, \\ u - g &> 1.88(g - r) + 0.68. \end{aligned} \quad (4)$$

This selection window is shown with the red lines in the middle panel of Fig. 10.

We combine our u^* -dropout and u -dropout number counts and plot them in the right-hand panel of Fig. 10. Our results (red symbols with error bars) are comparable to previous results in the CFHTLS-

Deep (black crosses, Hildebrandt et al. 2009a) at intermediate magnitudes ($r \sim 23 - 25$). However, at fainter magnitudes it is clear that our sample goes deeper than this previous work; it will go deeper still once the HSC-SSP imaging reaches its full depth at the conclusion of the survey. It is also worth emphasizing that our sample covers significantly more area than the $\sim 4\text{deg}^2$ CFHTLS-Deep work of Hildebrandt et al. (2009a). We will present $z \sim 3$ LBG luminosity and correlation functions in forthcoming papers (C. Liu et al., in preparation and Y. Harikane et al., in preparation).

5 OUTLOOK

The CLAUDS U -band data are all in hand and combined with the existing HSC-SSP $grizy$ data set. We are now using these data to tackle a number of scientific studies, and several papers are in preparation.

While the CLAUDS U -band data acquisition is complete, the HSC-SSP project is ongoing, with more exposures being acquired in the fields covered by CLAUDS. We plan to regularly update our catalogues to reflect this growing $grizy$ depth and expect to publicly release our data set – both catalogues and images – in 2020.

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During the course of this work, we have made use of the cosmological calculator of Wright (2006).

REFERENCES

- Abell P. A. et al., 2009, preprint ([arXiv:0912.0201](https://arxiv.org/abs/0912.0201))
- Adelberger K. L., Steidel C. C., 2005, *ApJ*, 630, 50
- Aihara H. et al., 2018a, *PASJ*, 70, S1
- Aihara H. et al., 2018b, *PASJ*, 70, S2
- Aihara H. et al., 2019, preprint ([arXiv:1905.12221](https://arxiv.org/abs/1905.12221))
- Arcila-Osejo L., Sawicki M., 2013, *MNRAS*, 435, 845
- Arnouts S. et al., 2002, *MNRAS*, 329, 355
- Bassett R. et al., 2019, *MNRAS*, 483, 5223
- Bauman S. E. et al., 2014, in Peck A. B., Benn C. R., Seaman R. L., eds, Proc. SPIE Conf. Ser. Vol. 9149, Observatory Operations: Strategies, Processes, and Systems V. SPIE, Bellingham, p. 91491K
- Bertin E., Arnouts S., 1996, *A&AS*, 117, 393
- Bertin E., Mellier Y., Radovich M., Missonnier G., Didelon P., Morin B., 2002, in Bohlender D. A., Durand D., Handley T. H., eds, Astronomical Data Analysis Software and Systems XI. ASP, San Francisco, p. 228
- Betoule M. et al., 2013, *A&A*, 552, A124
- Bohlin R. C., Gordon K. D., Tremblay P.-E., 2014, *PASP*, 126, 711
- Bosch J. et al., 2018, *PASJ*, 70, S5
- Boulade O. et al., 2003, in Iye M., Moorwood A. F. M., eds, Proc. SPIE Conf. Ser. Vol. 4841, Instrument Design and Performance for Optical/Infrared Ground-based Telescopes. SPIE, Bellingham, p. 72
- Bradshaw E. J. et al., 2013, *MNRAS*, 433, 194
- Bridge C. R., Carlberg R. G., Sullivan M., 2010, *ApJ*, 709, 1067
- Broadhurst T. J., Taylor A. N., Peacock J. A., 1995, *ApJ*, 438, 49
- Brown A. G. A. et al., 2016, *A&A*, 595, A2
- Capak P. et al., 2004, *AJ*, 127, 180
- Comparat J. et al., 2015, *A&A*, 575, A40
- Connolly A. J., Csabai I., Szalay A. S., Koo D. C., Kron R. G., Munn J. A., 1995, *AJ*, 110, 2655
- Coupon J. et al., 2012, *A&A*, 542, A5
- Coupon J., Czapon N., Bosch J., Komiyama Y., Medezinski E., Miyazaki S., Oguri M., 2018, *PASJ*, 70, S7
- Cuillandre J.-C., Magnier E. A., Isani S., Sabin D., Knight W., Kras S., Lai K., 2002, in Qinn P. J., ed., Proc. SPIE Conf. Ser. Vol. 4844, Astronomical Telescopes and Instrumentation. SPIE, Bellingham, p. 501
- de Jong J. T. A. et al., 2017, *A&A*, 604, A134
- Ferrarese L. et al., 2012, *ApJS*, 200, 4
- Flaugher B. et al., 2015, *AJ*, 150, 150
- Fontanot F., Cristiani S., Monaco P., Nonino M., Vanzella E., Brandt W. N., Grazian A., Mao J., 2007, *A&A*, 461, 39
- Foucaud S., Conselice C. J., Hartley W. G., Lane K. P., Bamford S. P., Almaini O., Bundy K., 2010, *MNRAS*, 406, 147
- Grazian A. et al., 2009, *A&A*, 505, 1041
- Grogin N. A. et al., 2011, *ApJS*, 197, 35
- Guhathakurta P., Tyson J. A., Majewski S. R., 1990, *ApJ*, 357, L9
- Gwyn S. D. J., 2008, *PASP*, 120, 212
- Gwyn S. D. J., Hartwick F. D. A., 1996, *ApJ*, 468, L77
- Hildebrandt H., Pielorz J., Erben T., van Waerbeke L., Simon P., Capak P., 2009a, *A&A*, 498, 725
- Hildebrandt H., van Waerbeke L., Erben T., 2009b, *A&A*, 507, 683
- Hudelot P. et al., 2012, VizieR Online Data Catalog. p. II/317
- Ibata R. A. et al., 2017, *ApJ*, 848, 128
- Ivezić Z. et al., 2019, *ApJ*, 873, 111
- Iwata I. et al., 2009, *ApJ*, 692, 1287
- Kawanomoto S. et al., 2018, *PASJ*, 70, 66
- Konno A. et al., 2018, *PASJ*, 70, S16
- Kriek M. et al., 2015, *ApJS*, 218, 15
- Laigle C. et al., 2016, *ApJS*, 224, 24
- Le Fèvre O. et al., 2013, *A&A*, 559, A14
- Lilly S. J. et al., 2007, *ApJS*, 172, 70
- Magnier E. A., Cuillandre J.-C., 2004, *PASP*, 116, 449
- Magnier E. A. et al., 2016, preprint (astro-ph/1612.05242)
- Malavasi N., Pozzetti L., Cucciati O., Bardelli S., Cimatti A., 2016, *A&A*, 585, A116
- Masters D. C., Stern D. K., Cohen J. G., Capak P. L., Rhodes J. D., Castander F. J., Paltani S., 2017, *ApJ*, 841, 111
- Matsuda Y. et al., 2011, *MNRAS*, 410, L13
- Matsuoka Y. et al., 2018, *PASJ*, 70, S35
- McConnachie A. et al., 2016, preprint (astro-ph/1606.00043)
- McLure R. J. et al., 2013, *MNRAS*, 428, 1088
- Miyazaki S. et al., 2018, *PASJ*, 70, S1
- Moutard T., Sawicki M., Arnouts S., Golob A., Malavasi N., Adami C., Coupon J., Ilbert O., 2018, *MNRAS*, 479, 2147
- Ono Y. et al., 2018, *PASJ*, 70, S10
- Ouchi M. et al., 2004, *ApJ*, 611, 685
- Pickles A., 1998, *PASP*, 110, 863
- Rovilos E., Georgantopoulos I., Tzanavaris P., Pracy M., Whiting M., Woods D., Goudis C., 2009, *A&A*, 502, 85
- Salim S. et al., 2009, *ApJ*, 700, 161
- Savoy J., Sawicki M., Thompson D., Sato T., 2011, *ApJ*, 737, 92
- Sawicki M., 2012a, *PASP*, 124, 1208
- Sawicki M., 2012b, *MNRAS*, 421, 2187
- Sawicki M., Thompson D., 2005, *ApJ*, 635, 100
- Sawicki M., Thompson D., 2006a, *ApJ*, 642, 653
- Sawicki M., Thompson D., 2006b, *ApJ*, 648, 299
- Sawicki M., Yee H. K. C., 1998, *AJ*, 115, 1329
- Sawicki M. J., Lin H., Yee H. K. C., 1997, *AJ*, 113, 1
- Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525
- Scodreggio M. et al., 2018, *A&A*, 609, A84
- Scoville N. et al., 2007, *ApJS*, 172, 1
- Silverman J. D. et al., 2015, *ApJS*, 220, 12
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Sorba R., Sawicki M., 2011, *PASP*, 123, 777
- Steidel C. C., Giavalisco M., Pettini M., Dickinson M., Adelberger K. L., 1996, *ApJ*, 462, L17
- Steidel C. C., Adelberger K. L., Dickinson M., Giavalisco M., Pettini M., Kellogg M., 1998, *ApJ*, 492, 428
- Steidel C. C., Adelberger K. L., Giavalisco M., Dickinson M., Pettini M., 1999, *ApJ*, 519, 1

- Steidel C. C., Adelberger K. L., Shapley A. E., Pettini M., Dickinson M., Giavalisco M., 2000, *ApJ*, 532, 170
- Steidel C. C., Shapley A. E., Pettini M., Adelberger K. L., Erb D. K., Reddy N. A., Hunt M. P., 2004, *ApJ*, 604, 534
- Strauss M. A. et al., 2002, *AJ*, 124, 1810
- Szalay A. S., Connolly A. J., Szokoly G. P., 1999, *AJ*, 117, 68
- Takada M. et al., 2014, *PASJ*, 66, R1
- Tanaka M. et al., 2018, *PASJ*, 70, S9
- Tasca L. A. M. et al., 2017, *A&A*, 600, A110
- Toshikawa J. et al., 2016, *ApJ*, 826, 114
- Tudorica A. et al., 2017, *A&A*, 608, A141
- Wright E. L., 2006, *PASP*, 118, 1711
- York D. G. et al., 2000, *AJ*, 120, 1579
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