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The magnetic early B-type stars I: magnetometry and rotation

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ABSTRACT

The rotational and magnetic properties of many magnetic hot stars are poorly characterized, therefore the Magnetism in Massive Stars and Binarity and Magnetic Interactions in various classes of Stars collaborations have collected extensive high-dispersion spectropolarimetric data sets of these targets. We present longitudinal magnetic field measurements $\langle B_z \rangle$ for 52 early B-type stars (B5–B0), with which we attempt to determine their rotational periods P_{rot} . Supplemented with high-resolution spectroscopy, low-resolution Dominion Astrophysical Observatory circular spectropolarimetry, and archival *Hipparcos* photometry, we determined P_{rot} for 10 stars, leaving only five stars for which P_{rot} could not be determined. Rotational ephemerides for 14 stars were refined via comparison of new to historical magnetic measurements. The distribution of P_{rot} is very similar to that observed for the cooler Ap/Bp stars. We also measured $v \sin i$ and v_{mac} for all stars. Comparison to non-magnetic stars shows that $v \sin i$ is much lower for magnetic stars, an expected consequence of magnetic braking. We also find evidence that v_{mac} is lower for magnetic stars. Least-squares deconvolution profiles extracted using single-element masks revealed widespread, systematic discrepancies in $\langle B_z \rangle$ between different elements: this effect is apparent only for chemically peculiar stars, suggesting it is a consequence of chemical spots. Sinusoidal fits to H line $\langle B_z \rangle$ measurements (which should be minimally affected by chemical spots), yielded evidence of surface magnetic fields more complex than simple dipoles in six stars for which this has not previously been reported; however, in all six cases, the second- and third-order amplitudes are small relative to the first-order (dipolar) amplitudes.

Key words: magnetic fields – stars: chemically peculiar – stars: early-type – stars: magnetic field – stars: massive – stars: rotation.

1 INTRODUCTION

Magnetic fields are detected in approximately 10 percent of stars with spectral types earlier than about A5, most prominently the chemically peculiar Ap/Bp stars and the Of?p stars (e.g. Grunhut et al. 2012b; Wade et al. 2016; Grunhut et al. 2017). In contrast

to the magnetic fields of cool stars, which tend to display complex surface topologies as well as magnetic activity cycles associated with dynamos, the magnetic fields of early-type stars tend to be strong (several kG), stable over timespans of up to decades, and topologically simple i.e. primarily dipolar (Donati & Landstreet 2009). Their striking observational properties, in combination with recent advances in the theoretical understanding of the behaviour of magnetic fields in stellar radiative zones (e.g. Braithwaite & Spruit 2004; Braithwaite 2008, 2009; Mathis & Zahn 2005; Duez, Braithwaite & Mathis 2010), strongly support the hypothesis that massive star magnetic fields are fossil fields, i.e. relics of an earlier stage in stellar evolution (e.g. Neiner et al. 2015; Alecian et al. 2017).

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[†] Based on observations obtained at the CFHT which is operated by the National Research Council of Canada, the Institut National des Sciences de l'Univers of the Centre National de la Recherche Scientifique of France, and the University of Hawaii; and at the La Silla Observatory, ESO Chile with the MPA 2.2 m telescope.

Until recently fossil fields were only known to exist in Ap stars and in a few He-weak and He-strong Bp stars: the first magnetic field detection in a chemically normal B-type star was for β Cep (Henrichs et al. 2000), while knowledge of Ap star magnetism dates to the beginning of stellar magnetometry (Babcock 1947), and the magnetic nature of He-weak and He-strong Bp stars has been known since the 1970s (Wolff & Wolff 1976; Landstreet & Borra 1978; Borra & Landstreet 1979). The advent of high-dispersion spectropolarimeters optimized for stellar magnetometry and mounted on 2- and 4-m class telescopes enabled the detection of fossil magnetic fields in a larger number of targets across a broader range of spectral types. The Magnetism in Massive Stars (MiMeS) survey systematically observed stars earlier than about spectral-type B5 (Wade et al. 2016). In addition to demonstrating that the Of?p stars are invariably magnetic (e.g. Martins et al. 2010; Wade et al. 2011, 2012, 2015; Grunhut et al. 2017), the MiMeS survey also detected magnetic fields in a large number of B-type stars. The MiMeS large programs (LPs) were followed up by the Binarity and Magnetic Interactions in various classes of Stars (BinaMICS) and the B-fields in OB stars (BOB) LPs, which respectively searched for magnetic fields in close binary systems containing at least one hot star, and for relatively weak magnetic fields (Alecian et al. 2015; Fossati et al. 2015b; Schöller et al. 2017).

Many hot, magnetic stars display variable emission in H Balmer, Paschen, and Brackett lines (e.g. Walborn 1974; Grunhut et al. 2012a; Oksala et al. 2015a), ultraviolet (UV) resonance lines (Barker et al. 1982; Smith & Groote 2001), hard and overluminous X-rays (Gagné et al. 1997; Oskinova et al. 2011; Nazé et al. 2014; Leto et al. 2017), and radio emission (Drake et al. 1987; Linsky, Drake & Bastian 1992; Chandra et al. 2015; Leto et al. 2017). These emissions originate in their circumstellar magnetospheres, which are classified as either dynamical or centrifugal magnetospheres depending on whether rotational support plays a negligible or dominant role in shaping the corotating, magnetically confined stellar wind (Petit et al. 2013). In dynamical magnetospheres, plasma returns to the photosphere on dynamical time-scales due to gravitational infall, thus they show emission only when the mass-loss rate is high enough to replenish the magnetosphere at the same rate that it is emptied due to gravity. Therefore, in general, the only stars with dynamical magnetospheres in emission are the magnetic O-type stars (Petit et al. 2013). In centrifugal magnetospheres, rotational support of the corotating plasma enables accumulation of plasma over longer time-scales; thus the magnetic B-type stars, which have weaker stellar winds, generally show emission only when they have centrifugal magnetospheres (Petit et al. 2013). While this classification scheme has been successful in explaining the qualitative differences in the emission properties of magnetic OB stars (Petit et al. 2013), our understanding of these systems is hampered by substantial uncertainty regarding the rotational and magnetic properties of magnetic early-type stars, as a large number of these stars are fairly recent magnetic detections. Precise characterization of rotational and magnetic properties is particularly important for the magnetic B-type stars, which due to their relatively weak winds should only show magnetospheric emission when rapid rotation enables the formation of a centrifugal magnetosphere (Townsend & Owocki 2005; ud-Doula, Townsend & Owocki 2006; Petit et al. 2013).

In this paper (hereafter Paper I), we present an extensive new data base of high-spectral resolution circular spectropolarimetry for the magnetic early B-type stars, with which (supplemented with low-resolution spectropolarimetry, high-resolution spectroscopy, and archival photometry) we attempt to determine their rotational pe-

riods, measure line broadening parameters, and characterize the properties of their longitudinal magnetic field curves. These measurements will form the empirical underpinning of the oblique rotator models and magnetospheric parameters, which will be presented in Paper II. Section 2 gives an overview of the observing programs and observations. Section 3 outlines the magnetic measurements, conducted using multi- and single-element least-squares deconvolution (LSD) profiles, and with H lines. In Section 4, line broadening parameters are measured and rotational periods determined. Section 5 examines the longitudinal magnetic field curves with the goal of identifying stars with significant contributions from higher order multipoles to their surface magnetic field topologies, and discusses the distributions of rotational properties in comparison to non-magnetic stars of similar spectral type, and to the magnetic stars of other spectral types. Conclusions and next steps are outlined in Section 6. Periodograms and longitudinal magnetic field curves of individual stars are given in Appendix A.

2 OBSERVATIONS

2.1 Sample selection

The sample consists of those magnetic main sequence B-type stars earlier than B5 identified by P13, for which sufficient spectroscopic and spectropolarimetric data are available to evaluate their rotational and magnetic properties. Three stars with reported spectral types later than B5 (HD 36526, HD 105382, and HD 125823) are also included, as their effective temperatures are above 15 kK. This is a consequence of surface chemical abundance peculiarities affecting spectral-type determinations, with He-weak stars in particular having, as their name implies, weaker He lines than expected for their effective temperatures. Since He lines are the primary diagnostic for spectral typing amongst hot stars, the spectral types assigned to He-weak stars are systematically later than would be implied by their effective temperatures. We include five additional stars, discovered to be magnetic since the P13 sample was published: HD 23478 (B3 IV, Sikora et al. 2015; Hubrig et al. 2015); the secondary star of the HD 136504 system (ϵ Lupi, B2 IV/V, Shultz et al. 2015b), in which the primary was already known to be magnetic (Hubrig et al. 2009; Shultz et al. 2012); HD 164492C (B1.5 V), HD 44743 (β CMa, B1 II/III), and HD 52089 (ϵ CMa, B1.5 II), discovered by the BOB collaboration (Hubrig et al. 2014; Fossati et al. 2015a). While HD 52089 was initially reported to be in the post-main-sequence phase of its evolution, analysis of its stellar parameters by Neiner et al. (2017a) has shown that it is likely still in the core hydrogen burning phase. In total, the initial sample consists of 52 early B-type stars reported to host magnetic fields.

The sample is summarized in Table 1. Stars are listed in order of their HD number; ALS 3694, which does not appear in the Henry Draper catalogue, is listed last. The second column gives alternate designations. The third column gives the spectral type and luminosity class. Remarks as to chemical peculiarity (He-weak or -strong), binarity (SB1/2/3), and/or pulsation (β Cep or Slowly Pulsating B-type star, SPB) are made in the fourth column. The remaining columns provide the number of spectropolarimetric, spectroscopic, and photometric observations available for each target.

As a comprehensive sample drawn from the literature, this study is neither volume nor magnitude limited. The statistical properties of the MiMeS survey were summarized by Wade et al. (2016). The MiMeS survey was complete for all OB stars up to $V \sim 1$, 50 per cent complete up to $V \sim 3$, and overall observed 7 per cent of the OB stars with $V < 8$. In spectral type, completeness was highest

Table 1. Summary of the sample stars and available spectropolarimetric, spectroscopic, and photometric observations. Stars are listed in order of HD number, with the exception of ALS 3694 which is not in the Henry Draper catalogue. The same observations were used for HD 136504A and HD136504B, therefore these are not counted twice in the table. Remarks indicate the type of chemical peculiarity (He-strong or He-weak), the star's binary status (SB1/2/3), and/or pulsational variability (β Cep or SPB). The next four columns give the number of high-dispersion spectropolarimetric observations obtained with ESPaDOnS (E), Narval (N), HARPSpol (H), and the total number (T) of magnetic measurements for each star, once low-S/N observations have been removed and spectral binning has been performed. The next column gives the number of DAO (D) observations obtained with the low-dispersion dimaPol spectropolarimeter. The number of spectroscopic observations obtained with FEROS (F) are given in the next column. The final column gives the number of archival *Hipparcos* (Hip) photometric measurements.

HD number	Star	Spec. type	Remarks	E	N	H	T	$\langle S/N \rangle$	D	F	Hip
3360	ζ Cas	B2 IV	SPB	–	56	–	56	1771	1	–	194
23478		B3 IV	He-s	14	–	–	10	555	11	–	74
25558	40 Tau	B3 V	SB2, SPB	12	19	–	31	994	–	–	90
35298		B3 Vw	He-w	10	2	–	12	541	16	–	119
35502		B5 V	SB3, He-w	4	20	–	24	486	24	33	98
35912	HR 1820	B2/3 V		12	–	–	6	317	3	–	130
36485	δ Ori C	B3 Vp	SB2, He-s	13	–	–	11	407	10	10	–
36526	V1099 Ori	B8 Vp	He-w	11	–	–	10	318	14	–	–
36982	LP Ori	B1.5 Vp	HeBe, He-s	18	5	–	14	471	–	–	–
37017	V 1046 Ori	B1.5-2.5 IV-Vp	SB2, He-s	10	–	–	8	854	–	33	95
37058	V 359 Ori	B3 VpC	He-w	15	–	–	10	597	3	–	–
37061	NU Ori	B0.5 V	SB2	24	–	–	11	1132	1	–	97
37479	σ Ori E	B2 Vp	He-s	2	16	–	18	670	–	36	–
37776	V 901 Ori	B2 Vp	He-s	13	13	–	26	598	3	–	103
44743	β CMa	B1 II/III	β Cep	–	–	11	5	700	–	–	108
46328	ξ^1 CMa	B1 III	β Cep	56	–	–	54	928	–	–	221
52089	ϵ CMa	B1.5 II	4	–	–	8	4	550	–	–	149
55522	HR 2718	B2 IV/V	He-s	9	–	–	9	596	–	–	224
58260		B3 Vp	He-s	9	–	–	9	438	–	–	125
61556	HR 2949	B5 V	He-w	41	–	–	22	1227	–	6	282
63425		B0.5 V	SB1?	11	–	–	11	648	–	–	111
64740	HR 3089	B1.5 Vp	He-s	4	–	13	17	346	–	21	99
66522		B2 III	He-s	1	–	4	5	296	–	–	117
66665		B0.5 V	–	5	16	–	21	414	–	–	49
66765		B1/B2 V	He-s	1	–	10	8	449	–	11	117
67621		B2 IV	He-w	1	–	8	6	497	–	–	133
96446	V 430 Car	B1 IVp/B2 Vp	β Cep, He-s	–	–	10	10	271	–	–	107
105382	HR 4618	B6 III	He-w	–	–	3	3	939	–	–	175
121743	ϕ Cen	B2 IV	β Cep, He-w	8	–	7	15	629	–	–	76
122451	β Cen	B1	SB2, β Cep	–	–	283	14	1802	–	–	119
125823	a Cen	B7 IIIp	He-w	17	–	–	17	708	–	–	101
127381	σ Lup	B1 V	He-s	20	–	–	20	919	–	–	158
130807	ρ Lup	B5	He-w, SB1?	2	–	10	12	564	–	–	56
136504	ϵ Lup	B2 IV-V	SB2, β Cep	91	–	–	14	2748	–	–	88
136504	ϵ Lup B	B2 IV-V	SB2, β Cep	–	–	–	–	–	–	–	–
142184	HR 5907	B2 V	He-s	27	–	–	27	1080	–	–	83
142990	V 913 Sco	B5 V	He-w	16	–	–	15	1176	–	–	111
149277		B2 IV/V	SB2, He-s	23	–	–	23	320	–	–	65
149438	τ Sco	B0.2 V	–	–	–	–	0	0	–	–	–
156324		B2 V	SB3, He-s	21	–	3	21	314	–	11	–
156424		B2 V	He-s, SB1/ β Cep	9	–	3	11	265	–	12	–
163472	V 2052 Oph	B1/B2 V	β Cep	–	13	–	13	631	–	–	66
164492C	EM* LkHA 123	B1.5 V	SB3	17	–	6	23	461	–	8	–
175362	Wolff star	B5 V	He-w	24	–	–	23	765	–	–	89
176582	HR 7185	B5 IV	He-w	–	44	–	43	865	19	–	112
182180	HR 7355	B2 Vn	He-s	4	–	–	4	1426	–	–	46
184927	V 1671 Cyg	B2 Vp	He-s	24	–	–	21	562	11	–	161
186205		B2 Vp	He-s	15	–	–	10	221	7	–	226
189775	HR 7651	B5 V	He-w	11	15	–	26	698	14	–	123
205021	β Cep	B1 IV	SB2, β Cep	–	21	–	20	1439	–	–	120
208057	16 Peg	B3 V	SPB	13	8	–	21	1143	–	–	88
	ALS 3694	B1	He-s	16	–	–	6	145	–	13	–

for the earliest stars (70 percent of O4 stars with $V < 8$), declining towards later spectral types. For the B-type stars, the sample is 30 percent complete at B0, diminishing to about 15 percent at B5: thus, the sample is most complete for the least common spectral types. However, the sample includes essentially all known and confirmed magnetic stars with spectral types between B5 and B0.

2.2 Observing programs

The majority of the observations used in this paper were acquired under the auspices of the MiMeS LPs at the 3.6 m Canada–France–Hawaii Telescope (CFHT) using ESPaDOnS (Echelle SpectroPolarimetric Device for the Observation of Stars), the 2-m Bernard Lyot Telescope (TBL) using Narval, and the European Southern Observatory (ESO) La Silla Observatory 3.6 m Telescope using HARPSpol. Five spectroscopic binaries in the sample (HD 35502, HD 136504, HD 149277, HD 156324, and HD 164492C) have also been observed by the BinaMICS LPs at CFHT and TBL. The remainder of the high-resolution spectropolarimetric data were acquired by various programs using ESPaDOnS, Narval, and HARPSpol at CFHT,¹ TBL, and the ESO La Silla 3.6 m telescope, as well some data collected via the BRITEpol LPs at CFHT, TBL, and ESO (Neiner et al. 2017b). The total numbers of ESPaDOnS, Narval, and HARPSpol observations for each star are given in the fifth to seventh columns of Table 1.

The methodology, observing strategy, instrumentation, and scope of the MiMeS LPs were described in detail by Wade et al. (2016). The BinaMICS LPs have largely adopted the strategies of the MiMeS LPs.

While the MiMeS LPs ended in 2012, only limited data were available for several magnetic stars identified by the Survey Component. Therefore, the Targeted Component was extended in two CFHT/ESPaDOnS PI programs in 2013 and 2014. In total, 1159 spectropolarimetric observations were gathered (614 ESPaDOnS, 248 Narval, and 297 HARPSpol). After removal of low signal-to-noise ratio (S/N) observations and, where appropriate and necessary, binning of individual spectra obtained in close temporal proximity, the high-resolution magnetic data set consists of 792 individual magnetic measurements with a mean of 15 measurements per star.

In addition to the high-resolution circular spectropolarimetry that forms the core of the data set, we have also utilized several supplementary data sets for certain stars. For 14 stars, we have obtained low-resolution circular spectropolarimetry with the dimaPol instrument at the 1.8 m Plaskett Telescope (137 observations). For 11 stars, we have high-resolution spectroscopy from FEROS at the MPG La Silla Observatory 2.2 m telescope, both gathered from the ESO archive, and (for eight stars) presented here for the first time.² Finally, we have downloaded epoch photometry from the *Hipparcos* archive, where available.

2.3 Spectropolarimetry

ESPaDOnS is a high-spectral resolution spectropolarimeter with a resolving power of $\lambda/\Delta\lambda \sim 65\,000$ at 500 nm and a spectral range of 370–1050 nm across 40 spectral orders (e.g. Donati 1993). Narval is an identical instrument in all important respects. ESPaDOnS and

Narval data were reduced using LIBRE-ESPRIT (Donati et al. 1997). HARPSpol has a greater spectral resolving power than ESPaDOnS and Narval ($\lambda/\Delta\lambda \sim 100\,000$), and a narrower spectral range, 378–691 nm, with a gap between 524 and 536 nm, across 71 spectral orders (Mayor et al. 2003; Piskunov et al. 2011). HARPSpol data were reduced using a modified version of the REDUCER pipeline (Piskunov & Valenti 2002; Makaganiuk et al. 2011). The excellent agreement between observations obtained with ESPaDOnS and Narval was demonstrated by Wade et al. (2016), who also describe the reduction of ESPaDOnS, Narval, and HARPSpol data in detail. The new data acquired after 2013 have been processed in the same fashion as described by Wade et al. (2016). The agreement between ESPaDOnS/Narval and HARPSpol is less well studied, but comparisons have been performed for the magnetic B1 star HD 164492C by Wade et al. (2017), and for the late-type Bp star HD 133880 by Kochukhov et al. (2017), both finding very good agreement between the instruments.

The detection and measurement of stellar magnetic fields requires data with a very high S/N. Therefore, in some cases spectra were removed from the analysis when the S/N was insufficient to obtain a meaningful measurement (typically, when the maximum S/N per spectral pixel $S/N_{\max} \leq 100$, below which the uncertainty in the longitudinal magnetic field is generally on the order of several kG). In other cases, if the time difference between observations could reasonably be expected to be small compared to either the known rotational period or the minimum rotational period inferred from the projected rotational velocity, spectra acquired on the same night were binned in order to increase the S/N (e.g. HD 136504, for which each measurement consists of at least four spectropolarimetric sequences). Especially, noisy observations were not included in such binning. The final number of measurements used for magnetic analysis is listed in the eighth column of Table 1. The ninth column gives the median peak S/N per spectral pixel in the final data set, $\langle S/N_{\max} \rangle$. The data quality is in general high, with a median S/N across all spectropolarimetric sequences of 612. The log of all spectropolarimetric observations is provided online.

Wade et al. (2016) noted that data collected with Narval during the late summers of 2011 and 2012 were affected by occasional, random loss of control of a Fresnel rhomb. While this issue cannot produce spurious magnetic signatures, the accuracy of magnetic measurements conducted with these spectra cannot be trusted. Therefore, spectra from the time windows given by Wade et al. were excluded from magnetic analysis. Only two observations within these epochs, of HD 176582 acquired on 2011 August 24 and HD 35502 on 2012 September 21, show obvious inconsistencies with the expected (B_z) variations. Another outlier, outside the windows given by Wade et al., was found for β Cep on 2009 October 04; this measurement was also discarded.

In addition to the high-spectral resolution spectropolarimetry, 137 low-spectral resolution measurements collected with the dimaPol spectropolarimeter mounted on the 1.8 m Dominion Astrophysical Observatory (DAO) Plaskett Telescope are available for 14 stars. The number of DAO observations are given in the 10th column of Table 1. dimaPol has a spectral resolution of approximately 10 000, covering a 25 nm region centred on the rest wavelength of the $H\beta$ line. Observations for HD 23478, HD 35502, HD 176582, and HD 184927 have been reported by Sikora et al. (2015), Sikora et al. (2016), Bohlender & Monin (2011), and Yakunin et al. (2015), respectively. The remainder are presented here for the first time. Unpublished measurements are also available online. The instrument and reduction pipeline are described in detail by Monin et al. (2012).

¹ Program codes CFHT 13BC012, CFHT 14AC010, CFHT 17AC16, PI: M. Shultz; and CFHT 14BC011, PI: J. Sikora.

² Program codes MPA 092.A-9018(A) and MPA LSO22-P95-007, PI: M. Shultz.

2.4 Spectroscopy

Eleven stars, mostly with optical emission originating in their magnetospheres were observed using the FEROS spectrograph at the MPG La Silla 2.2 m telescope. FEROS is a high-dispersion echelle spectrograph, with $\lambda/\Delta\lambda \sim 48000$ and a spectral range of 375–890 nm (Kaufer & Pasquini 1998). The data were reduced using the standard FEROS Data Reduction System MIDAS scripts.³ The number of FEROS spectra are summarized in the 11th column of Table 1.

2.5 Photometry

Hipparcos (High precision parallax collecting satellite) was an astrometric space telescope, whose mission lasted from 1989 to 1993. While the primary aim was to obtain high-precision trigonometric parallaxes, it also obtained photometry for a large number of stars. These data are available for 36 of the sample stars. As chemically peculiar stars exhibit photometric variability due to surface chemical abundance spots, in some cases rotational periods can be determined using *Hipparcos* photometry. Photometric variability may arise due to other physical mechanisms, e.g. pulsation, but in such cases the photometric and magnetic data should not phase coherently. These data were acquired from the online archive (Perryman et al. 1997; van Leeuwen 2007). The number of *Hipparcos* measurements is given in the final column of Table 1.

3 MAGNETOMETRY

Zeeman signatures are visible in the Stokes *V* profiles of individual spectral lines of many of the sample stars, as illustrated in Fig. 1 for the case the HD 58260 (top). This is particularly the case for sharp-lined stars with strong magnetic fields, as is true for HD 58260. However, many of the stars in the sample have much broader spectral lines in which Zeeman signatures are intrinsically more difficult to detect, such as HD 23478, which has the same spectral type as HD 58260 (Fig. 1, bottom). Both stars have longitudinal magnetic fields $|B_z| \sim 1$ kG (Bohlender et al. 1987; Sikora et al. 2015). In such cases, multiline analysis methods such as LSD are required in order to measure the star's magnetic field.

3.1 Least-squares deconvolution

We utilized the LSD procedure (Donati et al. 1997) in order to maximize the per-pixel S/N of the polarization profiles. In particular, LSD was performed using the iLSD package (Kochukhov, Makaganiuk & Piskunov 2010). Line lists were obtained from the Vienna Atomic Line Database⁴ (VALD3; Piskunov et al. 1995; Ryabchikova et al. 1997; Kupka et al. 1999, 2000) using ‘extract stellar’ requests, for effective temperatures from 15 to 30 kK in increments of 1 kK.

To ensure a homogeneous analysis, the line lists were automatically cleaned by removing: H Balmer and Paschen lines; He lines with strongly pressure-broadened wings; lines blended with H or He lines; lines in spectral regions that are often significantly contaminated by telluric features, depending on the observing conditions; and lines in the spectral region affected by instrumental ripples between 618 and 628 nm. When the spectropolarimetric data set included both HARPSpol and ESPaDOnS/Narval data, the line lists

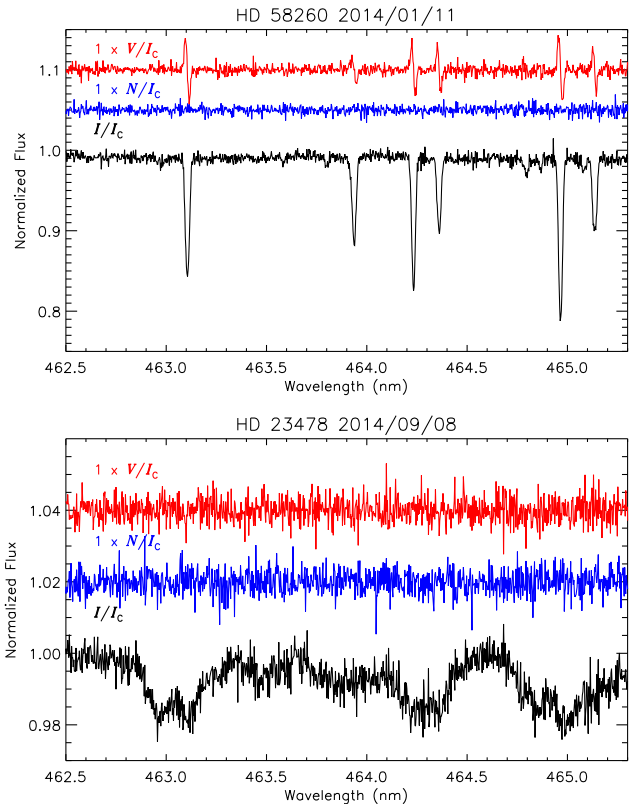


Figure 1. Stokes *I*, diagnostic null *N*, and Stokes *V* in a representative spectral window for the B3Vp star HD 58260 (top) and for the B3V star HD 23478 (bottom). Zeeman signatures are clearly detectable in individual spectral lines of HD 58260, due to the star's sharp spectral lines and strong magnetic field. However, despite the high S/N Zeeman signatures are not detectable in the individual lines of HD 23478, which has a similarly strong magnetic field but broad spectral lines. Compare the Stokes *V* profiles in this figure to the LSD Stokes *V* profiles for the same stars in Fig. 2.

were further truncated by removing all lines outside of the smaller HARPSpol spectral range. Line masks were then customized to each star by adjusting the depths of the remaining lines (typically 200–300 lines) to the observed line depths (e.g. Grunhut 2012). While He lines were by default excluded as their pressure-broadened wings do not have the same shape as metallic lines, and therefore do not conform to the fundamental LSD hypothesis, in many cases (in particular the He-fundamental stars), including some He lines can greatly improve the S/N of the LSD profile and, hence, the detectability of the Stokes *V* signature. Therefore, we also extracted LSD profiles with line masks including the He I lines 416.9, 443.8, 471.3, 501.6, 504.8, and 587.6 nm. Table 2 gives the number of lines used to extract LSD profiles for each star, for masks including only metallic lines (Z), and masks including both metallic and He lines (YZ). Measurements using H lines (X) are also provided; these are described in further detail in Section 3.2.2. For HD 44743 and HD 52089 (for which we did not extract new LSD profiles, but rely on the data published by Fossati et al. 2015a and Neiner et al. 2017a), Table 2 gives the number of lines used by Fossati et al. (2015a).

LSD profiles were extracted using wavelength, Landé factor, and line depth normalization constants of 500 nm, 1.2, and 0.1, respectively. Pixel velocity widths were set according to $v \sin i/40$ rounded to the nearest 1.8 km s^{-1} spectral pixel, with a minimum pixel width of 1.8 km s^{-1} adopted for all ESPaDOnS, Narval, and HARPSpol spectra, in order to increase the S/N in Stokes *V* in those stars

³ Available at: <https://www.eso.org/sci/facilities/lasilla/instruments/feros/tools/DRS.html>

⁴ Available at <http://vald.astro.uu.se/>

Table 2. Summary of magnetometry. For LSD profiles extracted with metallic lines (Z), metallic + He lines (YZ), and for measurements made using H lines (X), we provide the number of lines used in the mask (or the number of H lines to measure $\langle B_z \rangle$), the maximum $\langle B_z \rangle$ value obtained, and the mean significance level of the $\langle B_z \rangle$ measurements Σ_B (equation 1) made using that set of lines. The final three columns give: the number of single-element masks yielding $\Sigma_B > 3$ as a fraction of the total number of single-element masks with which LSD profiles were extracted; the significance A_e/σ_e of the elemental anomaly index (described in Section 3.2.1); and the set of measurements recommended for modelling.

Star	Number lines	Z $\langle B_z \rangle_{\text{max}}$ (kG)	Σ_B	Number lines	YZ $\langle B_z \rangle_{\text{max}}$ (kG)	Σ_B	Number of H lines	X $\langle B_z \rangle_{\text{max}}$ (kG)	Σ_B	N_3/N	A_e/σ_e	Sel.
HD 3360	253	-0.029 ± 0.007	1.4	271	-0.020 ± 0.004	1.2	3	-0.05 ± 0.05	0.9	0/13	2.2	YZ
HD 23478	140	-1.5 ± 0.2	4.4	163	-1.2 ± 0.1	8.0	2	-2.1 ± 0.2	7.9	1/10	2.9	X
HD 25558	–	–	–	224	-0.10 ± 0.04	1.3	–	–	–	0/0	–	YZ
HD 35298	176	-3.31 ± 0.10	16.8	189	-3.13 ± 0.09	17.1	3	3.2 ± 0.3	17.3	7/9	2.3	X
HD 35502	189	-2.6 ± 0.2	3.3	202	-2.6 ± 0.1	7.4	2	-1.58 ± 0.09	5.7	1/7	2.7	X
HD 36485	194	-2.57 ± 0.07	34.4	207	-2.36 ± 0.07	37.1	3	-2.17 ± 0.10	24.1	9/11	5.8	Z
HD 36526	179	3.2 ± 0.2	8.1	192	3.3 ± 0.3	4.9	3	3.4 ± 0.2	12.9	2/5	1.8	X
HD 36982	230	0.4 ± 0.8	1.9	272	0.34 ± 0.08	1.4	3	0.6 ± 0.2	1.6	0/9	1.8	YZ
HD 37017	234	-1.9 ± 2.1	2.9	253	-1.4 ± 0.3	3.7	1	-1.8 ± 0.3	7.3	1/6	1.3	X
HD 37058	166	-0.95 ± 0.03	17.7	179	-0.56 ± 0.03	9.4	3	-0.75 ± 0.05	7.1	7/11	4.6	X
HD 37061	284	-0.4 ± 0.2	1.5	326	-0.50 ± 0.08	2.5	1	-0.42 ± 0.08	2.6	0/9	0.9	YZ
HD 37479	222	-2.3 ± 0.3	5.3	236	2.0 ± 0.1	9.0	2	2.76 ± 0.09	13.6	4/10	4.1	X
HD 37776	253	-4.1 ± 0.1	15.2	268	-1.16 ± 0.07	5.1	2	1.7 ± 0.2	5.0	9/10	10.6	X
HD 44743	161	-0.023 ± 0.004	2.4	204	-0.026 ± 0.003	2.9	–	–	–	0/0	–	YZ
HD 46328	321	0.35 ± 0.01	54.4	343	0.31 ± 0.01	15.9	1	0.51 ± 0.06	8.0	8/10	1.9	Z
HD 52089	153	-0.011 ± 0.006	1.4	119	-0.018 ± 0.006	2.1	–	–	–	0/0	–	YZ
HD 55522	180	1.0 ± 0.1	4.1	192	-0.56 ± 0.10	4.0	3	0.9 ± 0.1	4.2	0/11	2.1	X
HD 58260	141	1.82 ± 0.03	66.3	142	1.77 ± 0.03	66.3	3	1.92 ± 0.06	21.7	8/9	4.5	Z
HD 61556	213	-0.89 ± 0.06	14.4	217	-0.89 ± 0.06	14.4	3	-0.82 ± 0.03	14.0	7/11	4.7	X
HD 63425	267	0.130 ± 0.006	9.8	283	0.101 ± 0.005	10.5	3	0.2 ± 0.1	1.1	3/10	1.0	Z
HD 64740	279	-0.9 ± 0.1	3.1	298	-0.68 ± 0.06	5.8	3	-0.82 ± 0.09	8.6	1/10	2.4	X
HD 66522	207	-0.54 ± 0.01	30.6	209	$-0.50 \pm 0.01s$	30.8	3	0.62 ± 0.07	10.2	8/11	7.0	X
HD 66665	284	-0.126 ± 0.009	4.6	299	-0.101 ± 0.009	4.0	3	0.24 ± 0.06	1.0	1/10	1.3	Z
HD 66765	243	-1.1 ± 0.1	6.8	288	-0.7 ± 0.2	12.9	3	-0.81 ± 0.03	15.7	3/10	2.0	X
HD 67621	250	0.42 ± 0.01	10.8	269	0.26 ± 0.01	8.8	3	0.29 ± 0.03	5.1	6/13	4.4	X
HD 96446	265	-1.00 ± 0.02	47.5	308	-1.20 ± 0.06	25.9	3	-1.40 ± 0.07	25.6	8/11	4.3	Z
HD 105382	160	-0.76 ± 0.03	14.2	173	-0.70 ± 0.03	11.4	3	-0.74 ± 0.03	16.3	4/9	2.1	X
HD 121743	225	0.39 ± 0.06	3.7	250	0.3 ± 0.1	5.0	3	0.33 ± 0.09	4.2	1/6	1.4	X
HD 122451	276	-0.03 ± 0.01	1.4	303	-0.065 ± 0.010	3.1	3	-0.05 ± 0.03	2.2	0/3	0.9	Z
HD 125823	230	-0.47 ± 0.01	22.9	243	0.30 ± 0.02	13.3	3	0.53 ± 0.07	7.4	6/11	4.0	X
HD 127381	269	0.17 ± 0.05	2.8	292	0.14 ± 0.02	2.8	3	-0.14 ± 0.10	1.6	0/10	2.2	Z
HD 130807	194	1.45 ± 0.03	23.8	207	1.00 ± 0.02	24.4	3	1.00 ± 0.03	17.1	4/10	9.9	X
HD 136504 A	193	-0.23 ± 0.02	9.5	329	-0.16 ± 0.01	16.3	–	–	–	0/0	–	Z
HD 136504 B	193	0.14 ± 0.05	2.6	329	0.10 ± 0.03	4.9	–	–	–	0/0	–	Z
HD 142184	148	-2.1 ± 0.4	2.3	161	-1.6 ± 0.1	7.9	2	-2.01 ± 0.10	14.8	1/8	2.7	X
HD 142990	193	1.11 ± 0.10	4.2	206	-0.9 ± 0.1	2.7	3	-1.2 ± 0.1	6.1	1/10	2.7	X
HD 149277	226	3.3 ± 0.1	7.7	239	2.34 ± 0.06	11.3	3	2.9 ± 0.2	5.2	6/9	3.7	X
HD 149438	267	0.088 ± 0.004	7.5	283	0.062 ± 0.005	5.5	2	0.18 ± 0.04	0.8	3/12	2.8	Z
HD 156324	218	2.7 ± 0.4	4.2	286	2.3 ± 0.1	9.3	2	2.2 ± 0.2	4.2	1/4	1.3	Z
HD 156424	233	-0.50 ± 0.05	8.9	238	-0.65 ± 0.03	15.3	3	0.27 ± 0.08	1.0	7/12	1.3	Z
HD 163472	266	-0.3 ± 0.1	1.2	306	0.09 ± 0.04	0.7	3	-0.19 ± 0.09	0.7	0/10	1.7	Z
HD 164492 C	307	2.8 ± 0.5	2.6	353	1.8 ± 0.1	6.5	1	2.3 ± 0.2	7.2	1/10	0.8	YZ
HD 175362	177	-5.19 ± 0.07	45.1	190	-4.44 ± 0.06	41.8	3	-5.13 ± 0.03	63.3	9/10	7.7	X
HD 176582	173	1.56 ± 0.09	9.3	186	1.40 ± 0.07	9.4	3	1.64 ± 0.06	17.4	4/6	2.0	X
HD 182180	144	2.4 ± 0.3	8.9	157	2.3 ± 0.1	16.4	2	2.7 ± 0.1	21.6	3/8	2.8	X
HD 184927	242	1.18 ± 0.03	22.5	255	0.65 ± 0.01	17.3	3	1.82 ± 0.07	25.7	10/13	14.1	X
HD 186205	213	-0.91 ± 0.04	21.8	233	-0.82 ± 0.03	28.2	3	-0.63 ± 0.04	7.9	9/10	2.8	YZ
HD 189775	152	1.21 ± 0.08	7.8	165	0.99 ± 0.06	11.0	3	1.29 ± 0.07	9.4	3/8	4.5	X
HD 205021	311	0.10 ± 0.01	4.2	344	0.076 ± 0.006	6.6	3	0.15 ± 0.03	2.0	3/9	1.3	YZ
HD 208057	196	0.2 ± 0.1	1.2	236	-0.13 ± 0.03	2.5	3	-0.23 ± 0.04	1.8	0/9	1.6	YZ
ALS 3694	141	0.9 ± 0.7	1.2	154	-1.1 ± 0.1	5.7	2	-3.7 ± 1.0	3.7	1/12	3.0	X

with especially broad spectral lines. This pixel width was chosen based on the average pixel velocity width of ESPaDOnS and Narval data. The velocity range used for deconvolution was $\pm 2v \sin i + v_{\text{sys}}$, where v_{sys} is the systemic or central velocity, ensuring inclusion of the full spectral line while minimizing contamination in spectra of

stars with narrow spectral features. For spectroscopic binary stars, velocity ranges were set using line width and radial velocity (RV) semi-amplitudes, which were determined from RV measurements performed using parametric fitting of individual line profiles using a proprietary IDL routine (Grunhut et al. 2017). Profiles were extracted

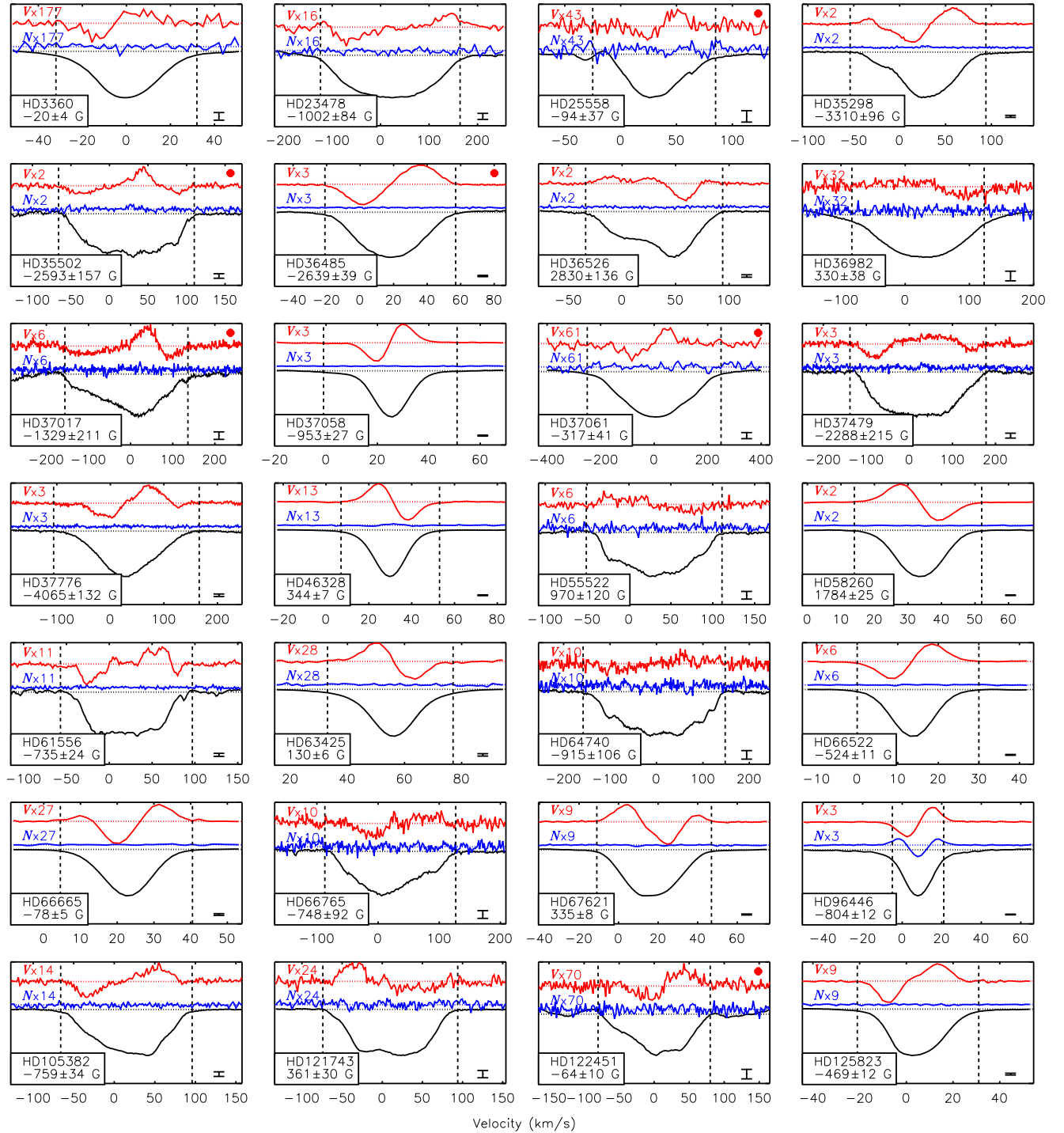


Figure 2. LSD profiles of individual stars. Stokes I in black, Stokes V in red, N in blue. The amplification factor applied to N and V is indicated in the top left of each panel. Vertical dashed lines indicate the integration ranges for determination of the FAP and $\langle B_z \rangle$. The mean error bar in N and Stokes V , scaled to the amplification factor, is displayed in the lower right. The legend on the lower left gives the name of the star, and $\langle B_z \rangle$: the LSD profiles shown here are those yielding the highest $\langle B_z \rangle$ significance. Stokes I profiles of binary stars corrected via disentangling are indicated by red circles in the upper right corner.

for Stokes I , V , and both diagnostic null N spectra. Representative LSD profiles for all stars except HD 44743, HD 52089 (for which we did not extract new LSD profiles), and HD 35912 (in which a magnetic field is not detected, see below) are shown in Figs 2 and 3.

As a first evaluation of the statistical significance of the magnetic signatures in the LSD Stokes V profiles, we calculated false alarm probabilities (FAPs) by comparing the signal inside the Stokes V

line profile to the signal in the wings (Donati, Semel & Rees 1992; Donati et al. 1997), where the boundaries of the line profile were determined from $\pm v \sin i$ adjusted to the rest frame of the star. Fig. 4 shows the cumulative distribution function of FAPs for all Stokes V profiles in the data set, where for each star we used the LSD profiles extracted with the mask yielding the highest S/N. 78 per cent of the LSD profiles have FAP $< 10^{-5}$, the upper boundary for a formal

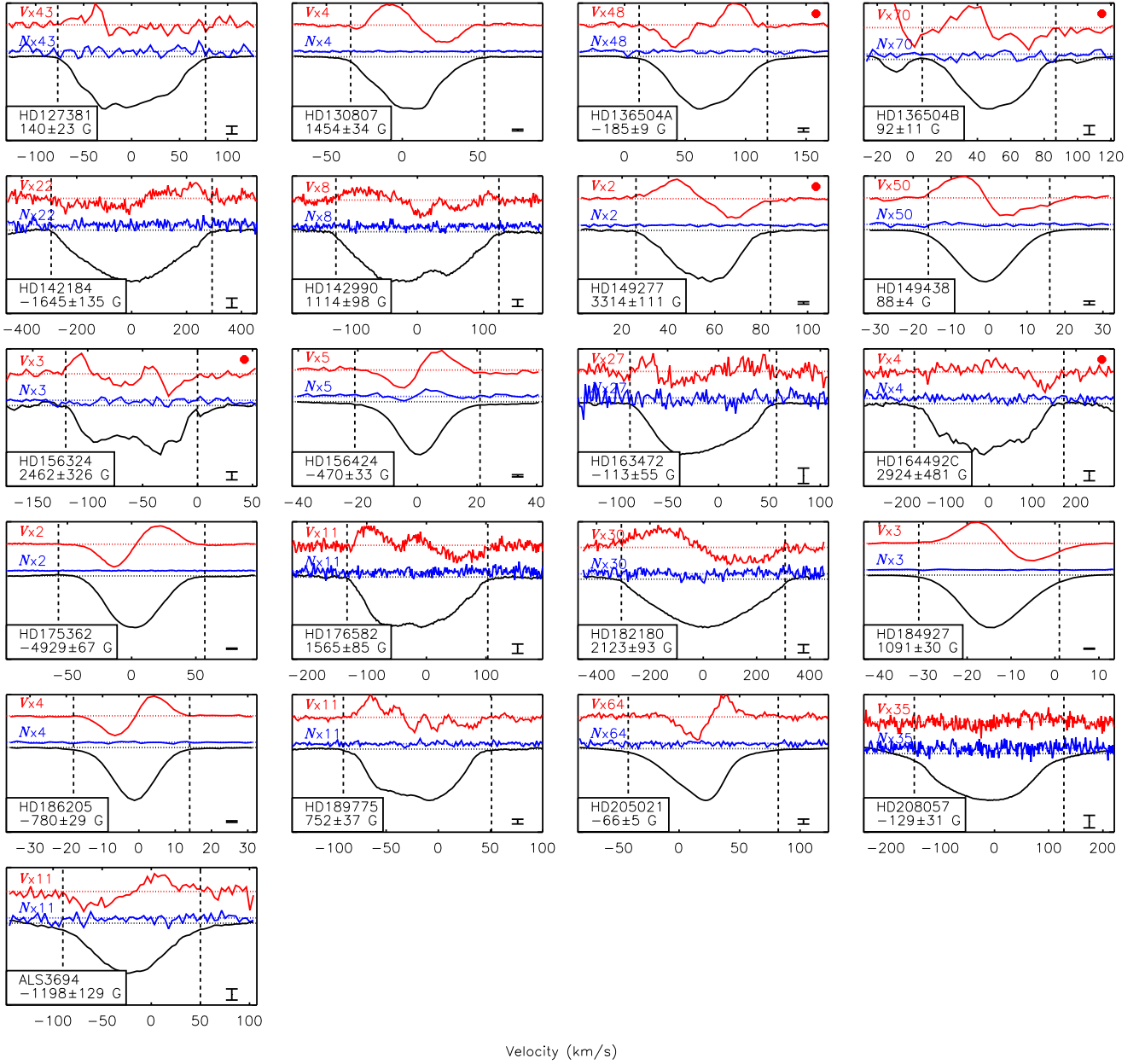


Figure 3. As Fig. 2. Note that the Stokes V LSD profiles of HD 135604 were not disentangled, and thus a contribution from the Zeeman signature of HD 136504A is apparent in the blue edge of HD 136504B's Stokes V profile.

definite detection (DD) (Donati et al. 1997). The remaining ~ 20 percent registering marginal detections with $10^{-5} < \text{FAP} < 10^{-3}$ or non-detections (NDs) with $\text{FAP} > 10^{-3}$ are accounted for by the presence in the data set of weak Zeeman signatures relative to the S/N.

Approximately 4 percent of N LSD profiles also register a DD (Fig. 4, blue line). This is due to stars exhibiting rapid RV variations over short timespans compared to the exposure times. Neiner et al. (2012b) noted the strong signal in the N profile of the β Cep star HD 96446, and we confirm its presence in all spectra in the data set (Fig. 2). Also showing DDs in some N spectra are HD 46328 (Fig. 2) and HD 205021 (Fig. 3), both β Cep stars; HD 156324, a short-period binary (Fig. 3); and HD 156424, which is not listed as either a binary or a β Cep variable, but does exhibit RV variations

(Fig. 3). The N signatures in HD 156324 and HD 156424 were noted by Alecian et al. (2014). Such N signatures can be reproduced by including RV variations between sub-exposures (Neiner et al. 2012b; Shultz et al. 2017). Correcting for rapid RV variations in the reduction stage does not result in changes to $\langle B_z \rangle$ larger than the error bars (Neiner et al. 2012b; Shultz et al. 2017). Therefore, RV variability should have no impact on $\langle B_z \rangle$.

ESPaDOnS measurements did not confirm a magnetic field in one star, HD 35912, catalogued as a magnetic star by Bychkov, Bychkova & Madej (2005) based on their own low-resolution measurements, as well as measurements performed by Conti (1970). The LSD profile deconvolved from the spectrum with the highest S/N is shown in Fig. 5. HD 35912 has sharp spectral lines, with $v \sin i = 12 \text{ km s}^{-1}$. All six of the validated observations, with a

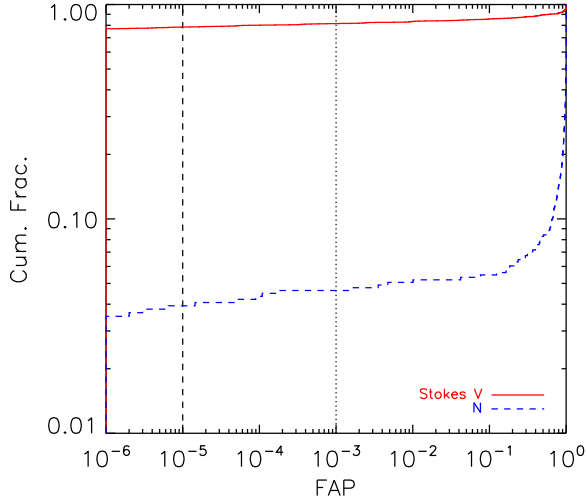


Figure 4. Cumulative distribution of FAPs. The dashed (dotted) vertical line indicates the lower cut-off for definite (marginal) detections. The ~ 80 per cent of Stokes V measurements are DDs, with $\text{FAP} \leq 10^{-5}$, while ~ 4 per cent of N measurements are also DDs.

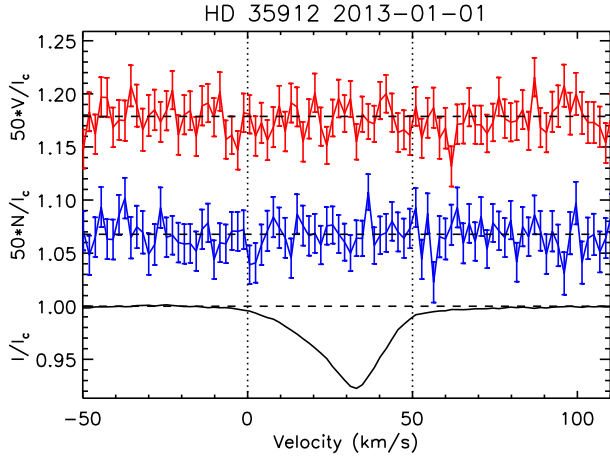


Figure 5. The highest S/N LSD profile of HD 35912. Dotted lines show the integration range used for calculation of FAP and measurement of $\langle B_z \rangle = -13 \pm 36$ G. The Stokes V profile (top) is indistinguishable from the N profile (middle).

median peak S/N per spectral pixel of 340, yielded NDs in the LSD profile. The median error bar of the longitudinal magnetic field measurements of this star is ~ 35 G, more than sufficient to detect the previously reported ≥ 6 kG magnetic dipole. DAO $\langle B_z \rangle$ measurements, with a mean error bar of 200 G, also failed to detect a magnetic field (Table 3). We therefore removed this star from the analysis.

3.2 Longitudinal magnetic field

To evaluate the strength of the magnetic field, we measured the line of sight or longitudinal magnetic field (B_z) in G, averaged over the stellar disc, by measuring the first-order moment of Stokes V normalized to the equivalent width (EW) of Stokes I (e.g. Mathys 1989). Integration ranges were set by $\pm v \sin i + v_{\text{sys}}$. The same measurement can be applied to the diagnostic N profile, yielding an analogous ‘null longitudinal magnetic field’ $\langle N_z \rangle$, with which $\langle B_z \rangle$

Table 3. Summary of DAO $\langle B_z \rangle$ measurements.

HD number	H β		Fe II 492.3 nm		He I 492.2 nm	
	$\langle B_z \rangle_{\text{max}}$ (kG)	Σ_B	$\langle B_z \rangle_{\text{max}}$ (kG)	Σ_B	$\langle B_z \rangle_{\text{max}}$ (kG)	Σ_B
3360	-0.17 ± 0.09	1.8	–	–	–	–
23478	-2.9 ± 0.4	4.8	–	–	–	–
35298	5.7 ± 0.5	5.3	-6.8 ± 0.5	4.8	–	–
35502	-4.0 ± 1.0	4.6	–	–	–	–
35912	-0.5 ± 0.2	1.6	–	–	–	–
36485	-3.1 ± 0.4	10	–	–	–	–
36526	4.0 ± 0.3	6.5	7 ± 2	3.2	–	–
37058	-0.6 ± 0.1	2.4	–	–	–	–
37061	-0.02 ± 0.26	0.06	–	–	–	–
37776	-3.9 ± 0.8	3.5	–	–	–	–
176582	-2.1 ± 0.2	7.6	–	–	–	–
184927	2.1 ± 0.3	6.0	–	–	1.5 ± 0.1	4.3
186205	-1.0 ± 0.2	2.1	–	–	–	–
189775	1.7 ± 0.3	3.3	1.1 ± 0.2	1.6	–	–

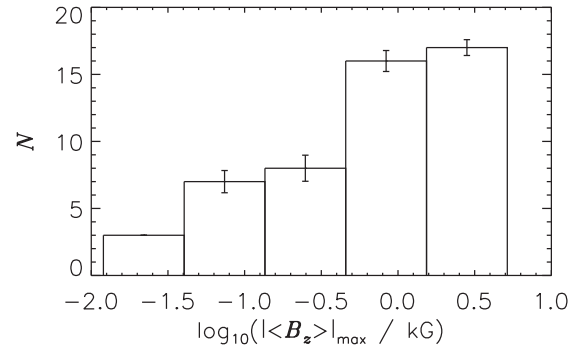


Figure 6. Histogram of maximum $|\langle B_z \rangle|$ values. Uncertainties were computed from the standard deviation in each bin via Monte Carlo simulations of synthetic data sets of the same size, with synthetic data points randomly shifted from the original within the measured error bars.

can be compared (Wade et al. 2000). The error bars σ_B and σ_N were obtained via error propagation of single-pixel photon noise uncertainties in the LSD profiles. Table 2 gives the value of the strongest $\langle B_z \rangle$ extremum, $\langle B_z \rangle_{\text{max}}$, in order to give an indication of the strength of the stars’ magnetic fields. These range from ~ 10 G (HD 52089) to ~ 5 kG (HD 175362). For some stars, the positive and negative $\langle B_z \rangle$ extrema are of a similar magnitude, leading to an apparent sign inversion when comparing $\langle B_z \rangle$ maxima obtained from different line lists. The histogram of $|\langle B_z \rangle|_{\text{max}}$ is shown in Fig. 6. The distribution peaks between 1.5 and 5 kG, with a cut-off above, and a low $\langle B_z \rangle$ tail extending to 10 G. Since strong magnetic fields are intrinsically easier to detect, the high- $\langle B_z \rangle$ cut-off likely reflects the intrinsic rarity of early B-type stars with higher values of $\langle B_z \rangle$. It is not so obvious that the distribution at low values reflects a real rarity of stars with relatively weak magnetic fields, as such fields are much harder to detect (see also the discussion of this possibility by Fossati et al. 2015a).

To evaluate the data quality, we calculated the mean significance Σ_B of the $\langle B_z \rangle$ measurements, which we define as:

$$\Sigma_B = \frac{1}{n} \sum_{i=1}^n \frac{|\langle B_z \rangle_i|}{\sigma_{B,i}}, \quad (1)$$

where i denotes an individual measurement out of the total number n of observations for a given star. For an individual measurement,

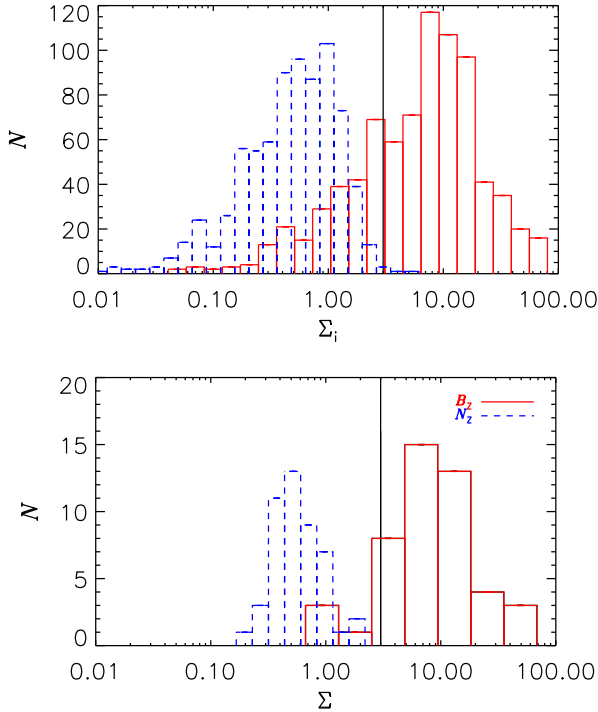


Figure 7. Histograms of $\Sigma_i = |\langle B_z \rangle|/\sigma_B$ and $|\langle N_z \rangle|/\sigma_B$ for all individual measurements (above), and the mean value of Σ for individual stars (below). The vertical solid lines indicate 3σ . The $\langle B_z \rangle$ Σ distribution peaks at 10σ , while the $\langle N_z \rangle$ distribution peaks below 1σ . The 67 per cent of $\langle B_z \rangle$ measurements are above 3σ , and only 0.2 per cent of $\langle N_z \rangle$ measurements. There are no stars for which the mean Σ_N is above 3σ significance, while 83 per cent of stars have Σ_B above the 3σ threshold.

$|\langle B_z \rangle|/\sigma_B$ is a measurement of the significance of the $\langle B_z \rangle$ measurement. Equation (1) then gives the mean of this significance across all measurements, with $\Sigma_B \leq 1$ indicating that the data set is dominated by noise. Σ_B is given in Table 2 for LSD profiles extracted using only metallic lines, and metallic plus He lines. For HD 44743 and HD 52089, Σ_B was evaluated from the $\langle B_z \rangle$ measurements presented by Fossati et al. (2015a) and Neiner et al. (2017a). Fig. 7 shows a histogram of these values, along with the corresponding histogram for Σ_N , the analogue of Σ_B for $\langle N_z \rangle$. The significance of $\langle B_z \rangle$ peaks at $\sim 10\sigma$, with a tail extending to 70σ . Approximately 10 per cent of the sample has $\Sigma_B \leq 3$: these are all stars with weak magnetic fields ($\langle B_z \rangle_{\max} \leq 300$ G). $\langle N_z \rangle$ is tightly clustered around 0.5σ , and is below 3σ for all stars. The DDs obtained from N profiles, examined in the previous sub-section, do not yield spurious signals in $\langle N_z \rangle$.

For DAO data, the longitudinal magnetic field is measured via the Zeeman shift between two spectra of opposite circular polarizations in the core of the $H\beta$ line, as well as nearby lines such as He I 492.2 nm or Fe II 492.3 nm. $\langle B_z \rangle$ is proportional to the Zeeman shift with a per-pixel scaling factor of 6.8 kG in $H\beta$, 6.6 kG in He I, and 14.2 kG in Fe II. The DAO $\langle B_z \rangle$ measurements are summarized in Table 3.

Ten of the sample stars are multilined spectroscopic binaries. The LSD profiles of these stars are indicated in Figs 2 and 3 with red circles in the top right corner. Since $\langle B_z \rangle$ is calculated using the centre of gravity of the Stokes I and V profiles, if the components are blended (as is generally the case), $\langle B_z \rangle$ can be affected by the contribution of binary companions to the Stokes I spectrum. In order to remove this influence, disentangled LSD profiles were obtained

via an iterative algorithm similar to that described by González & Levato (2006), and $\langle B_z \rangle$ was measured from the resulting Stokes I profiles of the magnetic component. Line profile modelling, RV measurement, and disentangling of Stokes I will be discussed in detail in the context of an analysis of the binary sub-population by Shultz et al. (in preparation). This correction is negligible for HD 36485 and HD 37061, due to the insignificant contribution of the non-magnetic star to Stokes I . For HD 35502, $\langle B_z \rangle$ is increased by ~ 25 per cent after correction. For HD 149277, the correction is non-existent for most observations, as the RV amplitude is much larger than the line widths and thus the line profiles are blended in only a few observations; however, for blended observations the correction for this star can be up to 40 per cent. For HD 122451 (in which the secondary is the magnetic star) and HD 156324, the correction is important. In the former case, $\langle B_z \rangle$ is approximately twice as high when measured using disentangled line profiles, since the components are blended in all observations and contribute approximately equally to the line profile. In the latter case, while the contributions of the non-magnetic stars are not large compared to the magnetic component, scatter in $\langle B_z \rangle$ is greatly reduced due to the strongly variable blending.

Iterative disentangling assumes that the Zeeman signatures in Stokes V are entirely due to one star. This assumption does not hold for HD 136504, in which both components are magnetic (Shultz et al. 2015b). Therefore, for this star the only measurements used were those obtained when the separation of the stellar line profiles in velocity space was greater than the summed $v \sin i$ of the components. This left only 5/14 observations available for analysis. Iterative disentangling was also not adopted for HD 25558: as both stars are SPB stars (Sódor et al. 2014), line profile variability in both components makes spectral disentangling unreliable. Therefore, we limited the HD 25558 data set to only those observations in which the line profiles are separately distinguishable (leaving 11/31 measurements), and used model fits to remove the flux of the non-magnetic primary using the binary line-profile fitting program described by Grunhut et al. (2017). As the EW ratio is a free parameter in model fits, this method is able to compensate for the changing EWs of the two components due to pulsations.

3.2.1 Measurements with different elements

While the relative precision of $\langle B_z \rangle$ is quite high, as evaluated by Σ_B , there is also the question of accuracy. Borra & Landstreet (1977) showed that $\langle B_z \rangle$ measurements of Ap stars displayed systematic differences when measured using spectral lines from different elements. This phenomenon has subsequently been reported for some Bp stars (Bychkov et al. 2005; Yakunin et al. 2011, 2015; Shultz et al. 2015a). The physical origin of these discrepancies is thought to be the same as that leading to the photometric variability of Ap/Bp stars, namely, surface chemical abundance spots. Spots lead to differential line formation such that there is a greater contribution to Stokes V in regions of enhanced abundance, causing the polarized flux to be enhanced in some regions relative to others and, hence, warping the Stokes V profile (Yakunin et al. 2015).

To explore the prevalence and influence of this effect in our sample, we extracted LSD profiles using single-element line masks, i.e. line masks in which lines of only a single chemical element were included. These were obtained from the cleaned and tweaked masks described in Section 3.1, with the criterion that a given element have at least three isolated lines in the analysis region. Table 2 gives the fraction of masks N_3/N , where N is the total number of masks for

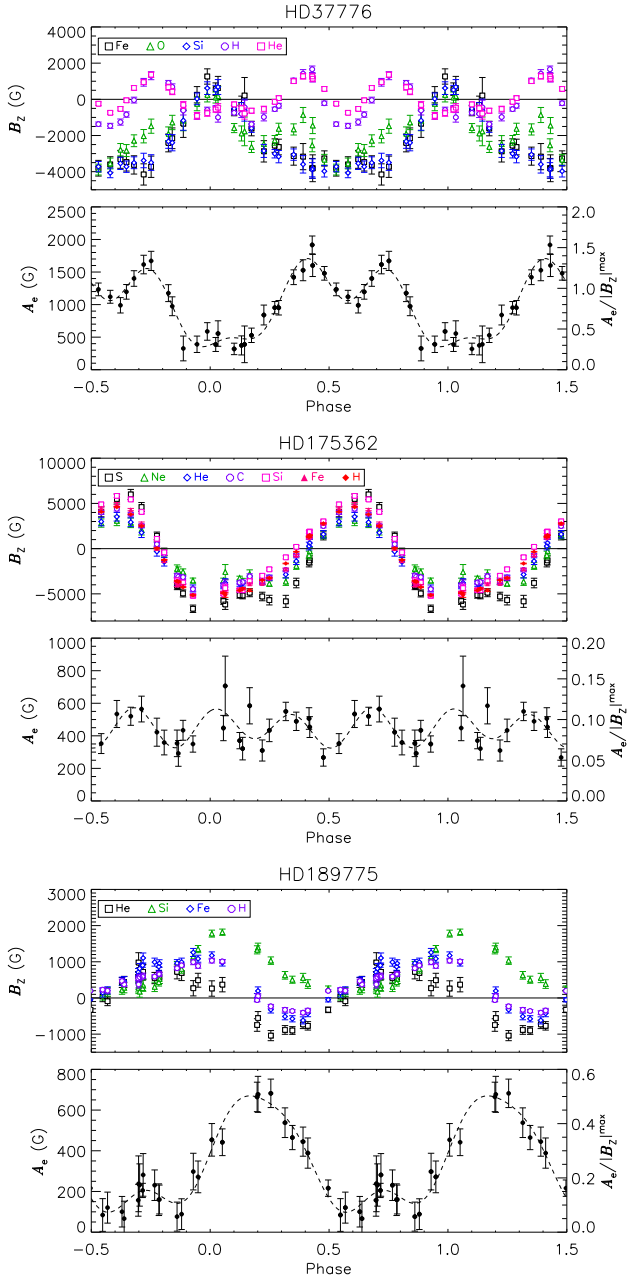


Figure 8. Variation in $\langle B_z \rangle$ when measured using different chemical elements for HD 37776, HD 175362, and HD 189775. Top panels: single-element (B_z) measurements. Bottom panels: elemental anomaly index A_e (equation 2). The left-hand axis gives the absolute A_e in G, the right-hand axis the relative A_e obtained via normalization to $|\langle B_z \rangle|_{\max}$. The dashed line indicates a third-order sinusoidal fit, used to determine the integrated value of A_e across all rotational phases (see the text). Note the variation between stars in both the magnitude and variability of A_e .

which LSD profiles could be extracted for each star, and N_3 is the number of masks for which $\Sigma_B > 3$. We also measured $\langle B_z \rangle$ using H lines (see below, Section 3.2.2). $\langle B_z \rangle$ measurements using different elements are shown in Fig. 8 for the examples of HD 37776, HD 175362, and HD 189775, phased according to the ephemerides given below in Section 4.2. All three stars show a large variation in $\langle B_z \rangle$.

To quantify the degree to which $\langle B_z \rangle$ differs when measured using the spectral lines of different elements, we calculated the *elemental*

anomaly A_e . This is the weighted standard deviation across all n single-element measurements obtained from a given spectrum i :

$$A_e = \sqrt{\frac{1}{\sum_{i=1}^n (\sigma_{B,i} - \overline{\sigma_B})^{-2}} \sum_{i=1}^n \frac{(\langle B_{z,i} \rangle - \overline{\langle B_z \rangle})^2}{(\sigma_{B,i} - \overline{\sigma_B})^2}} \quad (2)$$

where $\sigma_{B,i}$ is the uncertainty in $\langle B_z \rangle$ for element i , $\overline{\sigma_B}$ is the mean uncertainty across all n elements, and $\overline{\langle B_z \rangle}$ is likewise the mean $\langle B_z \rangle$ over all n elements. An error-bar-weighted standard deviation is used because measurements obtained from different single-element line masks have systematic differences in uncertainty, due to the large differences in the number of lines available for different elements. Normalization to $|\langle B_z \rangle|_{\max}$ is then performed in order to determine the fractional variation in $\langle B_z \rangle$ across different elements, rather than the absolute variation, which will be higher for stars with intrinsically stronger magnetic fields. The uncertainty σ_e in A_e was calculated from the weighted mean error bar across all elements.

Examples of A_e curves are shown in the bottom panels of Fig. 8, with the left axes giving A_e in G, and the right axes giving the fractional A_e . The behaviour of A_e is different with each star. In the case of HD 37776, A_e changes very significantly with rotation phase, peaking near one or more of the magnetic extrema, and reaching a minimum near $\langle B_z \rangle = 0$. HD 189775 also shows large changes with phase, but the largest difference is seen near the magnetic equator ($\langle B_z \rangle \sim 0$) due to a phase shift between Si $\langle B_z \rangle$ measurements and other elements. By contrast, HD 175362 shows little variation of A_e with rotation. The peak values are also quite different: 150 per cent for HD 37776, 50 per cent for HD 189775, and only 10 per cent for HD 175362. Stars with highly variable A_e likely exhibit greater surface abundance anisotropies than stars in which A_e is less variable.

In order to obtain a single number with which to characterize the average strength of deviations between the magnetic curves derived for different elements for a given star, we used the IDL CURVEFIT routine to generate second- or third-order harmonic best fits, which were then integrated as functions of rotational phase. The significance of this integrated value is then A_e/σ_e , which is given for each star in the second-last column of Table 2, where σ_e is the mean uncertainty in the individual A_e measurements. If $A_e/\sigma_e < 1$, it can be concluded that differences in $\langle B_z \rangle$ between different elements are a consequence of noise rather than real variations. It is this significance in which we are interested, as it will help to decide whether accurate measurements are available using metallic lines, or whether H lines should be used instead. Fig. 9 shows the histogram of A_e/σ_e for all stars. For the full sample (top panel), the median significance is 2.7σ . For the sub-sample of He-weak and He-strong Bp stars, the median of the distribution is at 2.9σ (second panel from the top). Amongst pulsating β Cep and SPB stars (third panel from the top) and the remainder of the sample with no apparent peculiarities (bottom panel), only one star, HD 96446, has $A_e/\sigma_e \geq 3$: as HD 96446 is also a He-strong star, this likely reflects the star's chemical peculiarities. The medians of these distributions are 1.4σ and 1σ , respectively. While there are only eight non-chemically peculiar stars for which A_e could be measured, these results are consistent with an origin of the effect in distortion due to chemical spots, and confirm that when chemical peculiarities are not present, there is no significant difference in $\langle B_z \rangle$ when measured using different chemical elements. Even amongst the Bp stars, although there are several stars for which $A_e/\sigma_e \gg 3$, differences in $\langle B_z \rangle$ measured from different elements are negligible for many of the stars.

It should be noted that while the mathematical treatment of A_e is based upon *random* uncertainties, these variations are in fact

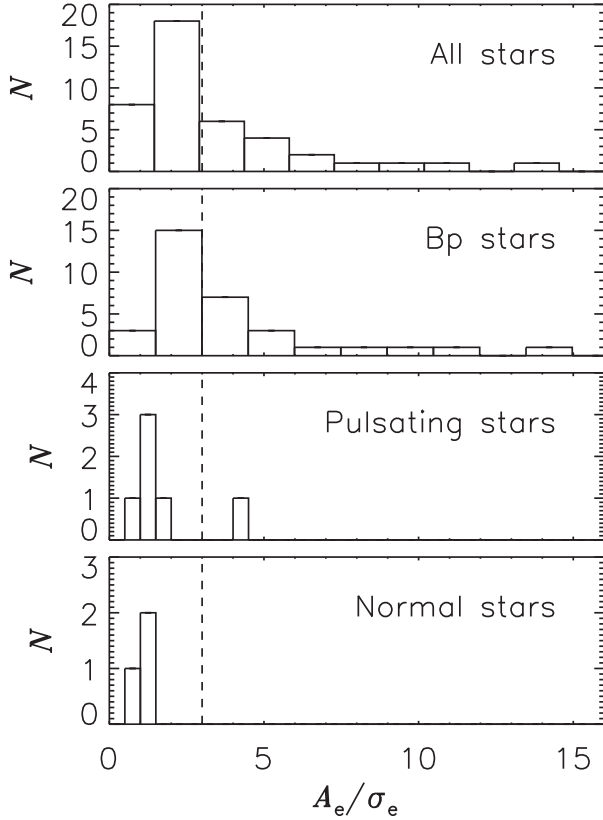


Figure 9. Significance histogram for A_e for (top–bottom) all stars in the sample for which A_e could be measured, Bp stars, pulsating stars, and chemically normal, non-pulsating stars. The dashed line marks 3σ significance. Approximately half the sample shows at least a 3σ spread in $\langle B_z \rangle$ measurements conducted with different elements. All of these are Bp stars: the only pulsator with $A_e/\sigma_e \geq 3$ is HD 96446, which is also a He-strong star.

systematic. Unfortunately, it is not obvious what simple mathematical tools might be utilized in order to capture the systematic variations between different sets of measurements. A rigorous exploration of this effect will require Zeeman Doppler Imaging (ZDI) of the surface magnetic fields together with Doppler Imaging (DI) of the chemical abundance patterns. As ZDI and DI cartographies are outside of the scope of this work, we have limited ourselves to determining A_e as a somewhat crude indicator for comparing the magnitude of the effect between different stars. However, it should be kept in mind that as of yet no pattern in the distribution of chemical elements relative to the surface magnetic field has been detected, hence, a treatment of these variations in terms of random error might not be entirely unwarranted.

3.2.2 H line $\langle B_z \rangle$ measurements

In addition to yielding different $\langle B_z \rangle$ measurements, chemical spots can also lead to anharmonic $\langle B_z \rangle$ variations that can be mistaken for contributions from higher order multipolar components to the photospheric magnetic field. Borra & Landstreet (1977) showed that measurements performed using H lines avoid this problem, as H is in general distributed relatively uniformly over the photosphere. Historically, such measurements were performed with photopolarimeters using the wings of the $H\beta$ line. However, Landstreet et al. (2015) have shown that, when high-resolution spectropolarimetry is available, $\langle B_z \rangle$ can be measured using the rotationally broadened

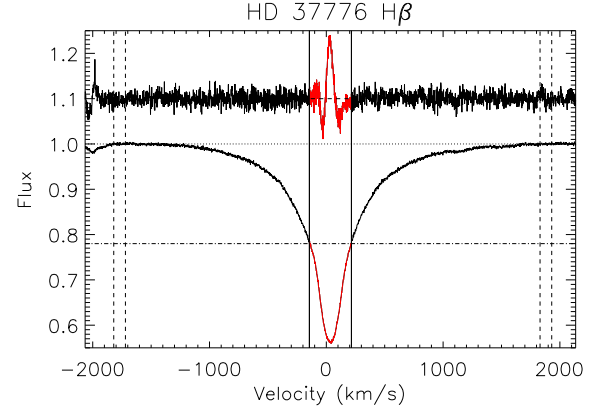


Figure 10. Measuring $\langle B_z \rangle$ from the $H\beta$ line of HD 37776. Stokes I (bottom) is normalized to the continuum. Normalization regions are shown by vertical dashed lines, with the ‘true continuum’ indicated by the horizontal dotted line. The integration limits are indicated by vertical solid lines. Note that the Stokes V (top) signature is entirely within the rotationally broadened non-LTE core of the line. The dot–dashed line indicates the ‘line continuum’, used to evaluate $\langle B_z \rangle$ (see the text).

non-local thermodynamic equilibrium (non-LTE) core of $H\alpha$ rather than using the line wings. Fig. 10 shows Stokes I and V for the $H\beta$ line of HD 37776, demonstrating that the magnetic signature is also clearly detectable in the Stokes V profile of this line.

We measured $\langle B_z \rangle$ using $H\alpha$ through $H\gamma$, except for stars with significant emission (as is the case for e.g. HD 37776, Petit et al. 2013), in which case $H\alpha$ measurements were not used. Lines at shorter wavelengths than $H\gamma$ were not used as the S/N of ES-PaDON/S/Narval spectra is typically much lower in this region. The final $\langle B_z \rangle$ was calculated as the error-bar-weighted mean across all Balmer lines used. Table 2 gives the number of H lines used for each star, as well as Σ_B for the H line measurements. We used the laboratory wavelengths of the lines and a Landé factor $g = 1$. The EW was measured using a continuum I_c taken at the boundary of the rotationally broadened line core, rather than the ‘true’ continuum bounding the pressure-broadened wings: this ensures the centroid of the line, and thus the separation of the circularly polarized components, is evaluated at the highest possible S/N and thus the maximum precision (Landstreet et al. 2015).

An example of the different result obtained using ‘true’ and ‘line’ continua is shown in Fig. 10 for HD 37776. If H-line $\langle B_z \rangle$ measurements are evaluated with I_c at the ‘true’ continuum rather than the boundaries of the non-LTE core (as determined by eye), the amplitude of the $\langle B_z \rangle$ curve is greatly reduced such that the measurements no longer agree with literature values, as demonstrated in Fig. 11 for HD 37776. We compare to photopolarimetric $H\beta$ wing measurements reported in the literature (Borra, Landstreet & Thompson 1983; Thompson & Landstreet 1985; Bohlender et al. 1987). Mathys et al. (2000) showed that, with H profiles calculated with a more accurate treatment of limb darkening and Stark broadening, photopolarimetric $\langle B_z \rangle$ measurements made using $H\beta$ lines should be corrected to ~ 80 per cent of their published values. The HD 37776 data have been phased with the non-linear ephemeris calculated by Mikulášek et al. (2008), which accounts for the spindown of the star. The modern data are more precise, but the general features of the $\langle B_z \rangle$ curve are essentially identical. Further comparisons of high-resolution H line $\langle B_z \rangle$ measurements with photopolarimetric data are shown in Appendix A for HD 36485, HD 37017, HD 37058, HD 64740, HD 125823, HD 142990, and HD 175362: in all cases, the agreement between modern and historical data is very

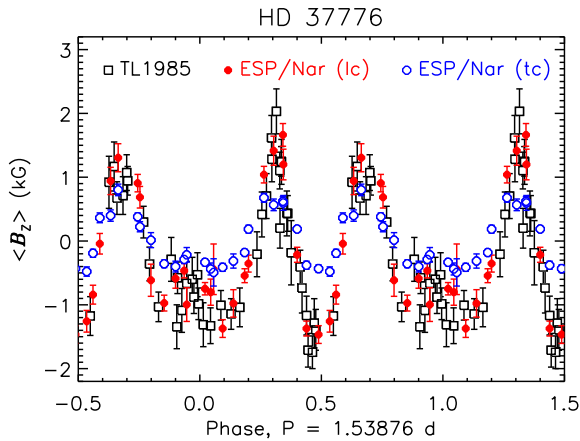


Figure 11. Photopolarimetric $H\beta$ wing measurements of HD 37776 reported by Thompson & Landstreet (1985) compared to ESPaDOnS and Narval data. The data are phased using the non-linear ephemeris determined by Mikulášek et al. (2008). The high-resolution spectropolarimetric measurements were renormalized either to the ‘true continuum’ (tc) or ‘line continuum’ (lc) (see the text and Fig. 10). The latter measurements provide a much better agreement with the historical data.

good. H line $\langle B_z \rangle$ measurements are also compared to $\langle B_z \rangle$ measurements from LSD profiles extracted from single-element line masks in Fig. 8: the reduced amplitude obtained when measuring $\langle B_z \rangle$ using the ‘true’ continuum would leave the amplitude of H line $\langle B_z \rangle$ measurements in disagreement with those obtained from other metallic lines.

We have also found that I_c should be evaluated in the same fashion for He lines with strong pressure-broadened wings. This introduces some ambiguity into $\langle B_z \rangle$ measurement for LSD profiles extracted using line masks dominated by He lines: to minimize this problem, He lines with broad wings were excluded from the line masks whenever possible.

A further consideration relates to the nature of echelle spectra. Before extracting LSD profiles from ESPaDOnS or Narval data, the spectra are in general normalized using a polynomial spline fit to the continuum of each echelle order. However, the broad wings of the H Balmer lines, especially $H\beta$ and $H\gamma$, overlap with the edges of their respective echelle orders: thus, polynomial normalization may distort the line profiles. This will then change the EW, and lead to an incorrect measurement of $\langle B_z \rangle$. To avoid this, we used spectra that had not been normalized using polynomial splines, instead normalizing using a linear fit between the edges of the line cores. Experimentation with different stars indicated that this strategy minimizes scatter in the measurements. As the REDUCE pipeline does not perform global normalization after merging the sub-exposures, HARPSpol data should not be affected by this issue.

4 ROTATION

4.1 Velocity broadening

Line-profile fitting was utilized to measure $v \sin i$. ESPaDOnS, Narval, and HARPSpol spectra combine a high-spectral resolution with a large spectral range, and offer numerous resolved metallic absorption lines with which to measure line broadening. In order to identify an optimal set of spectral lines, we first searched the VALD3 line lists described in Section 3.1 for isolated metallic lines. The final list for all stars includes: C II 426.7 and 658.2 nm; N II 404.4 and

568.0 nm; O II 418.5 and 445.2 nm; Ne I 640.2 nm; Ne II 439.2 nm; Si II 412.8, 504.1, 637.1, and 567.0 nm; Si III 455.3 and 457.5 nm; Si IV 411.6 nm; S II 543.3 and 566.5 nm; S III 425.4 nm;⁵ and Fe II 526.0 and 538.7 nm. For each star, the list was curated to remove lines that were absent (due to chemical peculiarities or effective temperature), or blended with other lines (due to high $v \sin i$). For HD 37061, for which none of the given lines were detected, we selected He I 501.6 nm (pressure broadening being fairly low in this line), He II 468.6 nm, and Si III 456.8 nm.

We used mean spectra created from all available spectra for each star, so as to minimize the impact of line profile variability. For binary stars, using the same method as in Section 3.2, we first decomposed the spectra into their stellar components using an iterative algorithm similar to that described by González & Levato (2006).

Line broadening was evaluated using a χ^2 goodness-of-fit test, comparing each line to a grid of synthetic line profiles covering a range of $v \sin i$ and v_{mac} values. The synthetic profiles were convolved with a Gaussian with a full width at half-maximum (FWHM) corresponding to the instrumental resolution of ESPaDOnS and Narval spectra, providing a lower limit to line broadening measurements of about 2 km s^{-1} . Examples of the χ^2 landscapes are shown in Fig. 12 for three stars: the sharp-lined star HD 63425; the He-weak star HD 35298, with an intermediate rotational period; and HD 37479, a rapidly rotating star.

Disc integration was performed with an engine similar to that described by Petit & Wade (2012), with some modifications. First, local profile widths were calculated using Maxwellian velocity distributions appropriate to the stellar T_{eff} and the atomic weight. Second, radial–tangential macroturbulence rather than isotropic turbulence was implemented (Gray 1975). This was motivated by the inclusion of higher mass stars in the sample, especially the pulsating β Cep stars. For Bp stars, v_{mac} values may be fictitious in that they do not likely reflect actual velocity fields within the stellar atmosphere, but can be taken as standing in for distortions to the line profile introduced by chemical spots or Zeeman splitting (see below). In general, macroturbulence may arise from a variety of physical processes, and can be taken as a stand-in for non-rotational broadening (Simón-Díaz et al. 2017).

For slow rotators (e.g. HD 63425, Fig. 12, top), solutions with high v_{mac} and $v \sin i \sim 0 \text{ km s}^{-1}$ produce much better fits. For intermediate rotators such as HD 35298 (Fig. 12, middle), the quality of the fit is improved by inclusion of non-zero v_{mac} , although the lowest χ^2 solutions with $v_{\text{mac}} = 0 \text{ km s}^{-1}$ yield essentially the same $v \sin i$ as the best-fitting solution with higher values of v_{mac} . For rapid rotators, inclusion of v_{mac} makes very little difference, as shown for the example of HD 37479 in Fig. 12 (bottom).

The final values of $v \sin i$ and v_{mac} were taken as the mean of the best-fitting values across all analysed spectral lines. Uncertainties were determined based on the standard deviation of the best-fitting parameters across all lines, and are typically on the order of 5 km s^{-1} . $v \sin i$ and v_{mac} are given in Table 4.

Sundqvist et al. (2013b) examined $v \sin i$ diagnostics for magnetic O-type stars known to have extremely long rotation periods, such that the true $v \sin i$ should be essentially zero, and found that in such cases $v \sin i$ was often drastically overestimated, by up to about 50 km s^{-1} (i.e. a similar value to that found for v_{mac}). They concluded that macroturbulence was likely contaminating

⁵ While this line is blended with an O II line, the strength of this O II line is negligible below about 20 kK; this line was not used for stars above this T_{eff} .

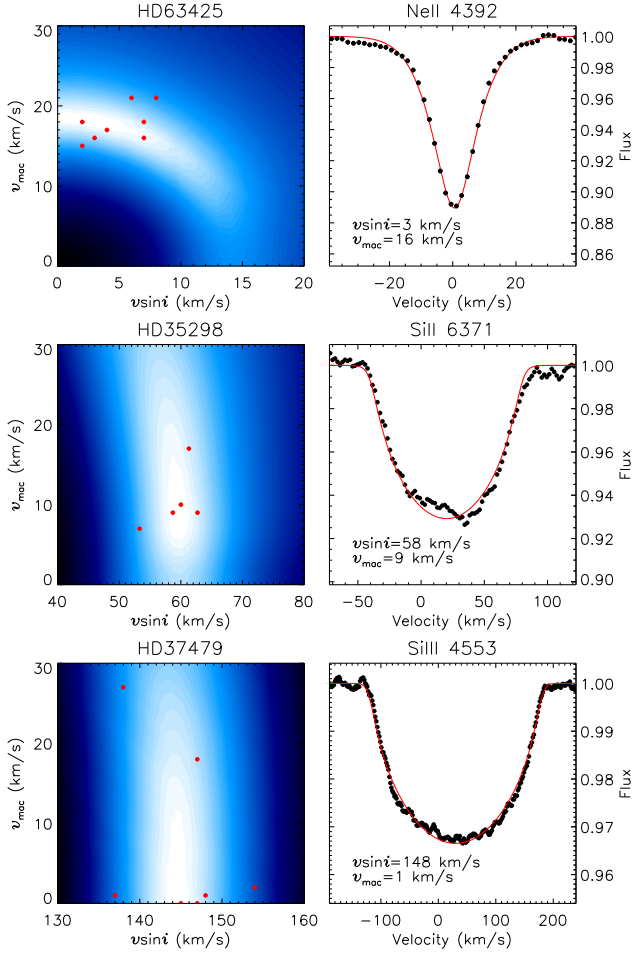


Figure 12. Left: $v \sin i$ versus $v_{\text{mac}} \chi^2$ landscapes. Shading is proportional to $\log(\chi^2)$, with lighter colours indicating lower χ^2 values. The map reflects the total χ^2 across all lines tested. The best-fitting models for individual lines are indicated by red circles. Right: representative model fits. Observed and synthetic line profiles are indicated by black circles and the red line. Top: the slow rotator HD 63425. Lines used for analysis were: (O II 418.5 and 445.2 nm; N II 567.9 nm; Ne II 439.2 nm; Al III 451.2 nm; Si III 455.3 and 457.5 nm; Si IV 411.6 nm). Middle: the intermediate rotator HD 35298. Due to strong blending in this He-weak Bp star, only five suitable lines could be found: Ne I 640.2 nm, Si II 412.8, 504.1, 634.7, and 637.1 nm. Bottom: the rapid rotator HD 37479. The lines used were C II 426.7 nm, N II 568.0 nm, O II 466.2 nm, Ne I 640.2 nm, Si II 637.1 nm, and Si III 455.3 and 456.8 nm.

the measurements. Aerts et al. (2014) found that measurements of $v \sin i$ could, at least in the case of the fast Fourier transform method, be significantly affected by both pulsation and chemical spots, as $v \sin i$ was strongly modulated with known pulsation and rotation periods. Both of these studies suggest that uncertainties in $v \sin i$ may in some cases be underestimated. Therefore, in case of stars for which we determine extremely long (>1 yr) rotation periods, we consider our $v \sin i$ measurements to be upper limits based upon the spectral resolution of the data, as the true projected rotational velocities are necessarily much lower, and as the measurements are similar in magnitude to the instrumental profile.

Another mechanism that may affect line broadening is gravity darkening. Townsend, Owocki & Howarth (2004) showed that, for stars with surface equatorial rotational velocities above 80 per cent

of their critical velocities, line broadening ceases to be a sensitive measure of the projected rotational velocity. This is because gravity darkening reduces the contribution of equatorial regions to the integrated flux. Accounting for this requires detailed spectral modelling including meridional temperature and surface gravity variations, together with knowledge of the inclination of the rotational axis from the line of sight. Only two stars in this sample are rotating in this regime, HD 142184 (Grunhut et al. 2012a), and HD 182180 (Rivinius et al. 2013). In both cases, careful spectral modelling accounting for oblateness as well as the meridional spectral differences arising from gravity darkening has been performed (Grunhut et al. 2012a; Rivinius et al. 2013), and the $v \sin i$ values so obtained are adopted here.

The line profiles of magnetic stars are subject to additional line broadening due to Zeeman splitting. For the majority of stars in the sample, this does not significantly affect line broadening: a 1 kG field will cause a line with an effective Landé factor $g_{\text{eff}} = 1$ to split by $\Delta\lambda \approx 0.0012$ nm at 500 nm, which is equivalent to about 0.7 km s^{-1} . However, for stars with strong surface magnetic fields (5–10 kG) and sharp spectral lines, Zeeman splitting can be a significant source of additional broadening as it is comparable to the broadening due to rotation and turbulence. Fig. 13 shows the predicted Zeeman broadening v_{Zeeman} as a function of the uncertainty in $v \sin i$, σ_v . The degree of Zeeman splitting was computed for a fictitious line with $\lambda_0 = 500$ nm and $g_{\text{eff}} = 1$, taking the surface strength of the magnetic field to be $3.4(B_z)_{\text{max}}$ (the minimum surface magnetic field strength at the magnetic pole, assuming a dipolar magnetic field). For the majority of the sample $\sigma_v \geq v_{\text{Zeeman}}$, and in these cases Zeeman splitting can be neglected. However, there are 11 stars for which $\sigma_v \leq v_{\text{Zeeman}}$. These are: HD 35298, HD 35502, HD 36485, HD 36526, HD 37776, HD 58260, HD 96446, HD 149277, HD 175362, HD 182180, and HD 184927.

For HD 36485, HD 58260, HD 96446, HD 149277, HD 175362, and HD 184927, the S II 566.5 nm line is detectable. This line has the lowest effective Landé factor in the VALD line lists, $g_{\text{eff}} = 0.5$, and is therefore much less strongly affected by Zeeman splitting. Fig. 13 also shows the amount of Zeeman splitting expected for S II 566.5 nm, for which v_{Zeeman} is much closer to σ_v . Therefore, for these stars, we used this line to measure line broadening, obtaining uncertainties from the $1\sigma \chi^2$ contours. For HD 35298 and HD 36526, the Fe II 450.8 nm line, which also has $g_{\text{eff}} = 0.5$, was used instead. In general, using these lines reduced $v \sin i$ by $3\text{--}5 \text{ km s}^{-1}$, consistent with the predicted degree of Zeeman splitting. For HD 37776, no low- g_{eff} lines could be detected, but the amount of Landé splitting ($\sim 10 \text{ km s}^{-1}$) predicted for this star according to the approximation described in the previous paragraph is close to the difference between the value determined above ($101 \pm 4 \text{ km s}^{-1}$), and the value reported by Kochukhov et al. (2011), 91 km s^{-1} , who performed a detailed spectrum synthesis accounting for the star's strong magnetic field, which we adopt here. For HD 35502 and HD 182180, the lines with $g_{\text{eff}} \leq 0.5$ are all too weak to detect, even in the mean disentangled spectrum, but $v_{\text{Zeeman}} \sim 6 \text{ km s}^{-1}$ is only slightly higher than $\sigma_v = 5 \text{ km s}^{-1}$, therefore we do not expect Zeeman splitting to significantly affect the $v \sin i$ measurements of these two stars. The $v \sin i$ values for these stars listed in Table 4 reflect the considerations described here.

Fig. 14 compares these values of $v \sin i$ to those collected by Petit et al. (2013). There is only one 3σ outlier, HD 35298, which is highlighted. The value given by P13, 260 km s^{-1} , is much higher than that found here, $58 \pm 2 \text{ km s}^{-1}$, as illustrated in Fig. 12 (middle). The original value appears to have been measured by McNamara

Table 4. Projected rotation velocities, macroturbulent velocities, and rotational periods. An asterisk next to the HD number indicates that the star's $\langle B_z \rangle$ curve is included in Appendix A. In most cases, JD0 (in HJD) is defined at $| \langle B_z \rangle |_{\max}$, as discussed in the text; ephemerides for which JD0 is defined otherwise are indicated by a superscript in the fifth column: maximum light: p ; minimum EW: s . The sixth column gives the method by which the period was determined: m : magnetic; u UV spectroscopy; s : optical spectroscopy; p : photometry; sp : a non-linear ephemeris accounting for the spin-down. Where more than one method was used to determine P_{rot} , the S/N in the seventh column is given in the same order as Method. The final column gives the reference for the period; where 'This work' is given together with another reference, we have refined the period determined in a previous study.

Star	$v \sin i$ (km s $^{-1}$)	v_{mac} (km s $^{-1}$)	P_{rot} (d)	JD0 −2400000 (d)	Method	S/N	Reference
HD 3360	19 ± 2	13 ± 1	5.37045(7)	45227.2(25) ^s	u	8.6	Neiner et al. (2003a)
HD 23478	136 ± 7	8 ± 8	1.0498(2)	47933.7(2) ^p	p	5.2	Sikora et al. (2015)
HD 25558*	35 ± 4	20 ± 10	1.233(1)	55400.0(2)	m	4.7	Sódor et al. (2014), This work
HD 35298*	58 ± 2	10 ± 8	1.85458(3)	54486.91(7)	p, m	8.0, 32.7	North (1984); Yakunin (2013), This work
HD 35502	78 ± 5	10 ± 9	0.853807(3)	56295.812850(3)	p, s, m	7.1, 20.7, 32.8	Sikora et al. (2016)
HD 36485*	26 ± 4	6 ± 5	1.47775(3)	48298.86(3) ^s	s	–	Leone et al. (2010)
HD 36526*	55 ± 5	5 ± 5	1.54185(4)	55611.93(6)	p, m	27.4	North (1984), This work
HD 36982*	86 ± 5	17 ± 12	1.8551(5)	54412.5(3)	m	5.9	This work
HD 37017*	134 ± 15	12 ± 22	0.901186(2)	43441.20(9)	m	11.0	Bohlender et al. (1987), This work
HD 37058*	11 ± 2	15 ± 2	14.581(2)	56522.0(4)	m	28.0	Pedersen (1979), This work
HD 37061*	189 ± 8	61 ± 30	1.0950(4)	55223.2(2)	m	11.0	This work
HD 37479	145 ± 5	13 ± 16	1.1908100(9)	42778.829(1) ^p	sp	–	Townsend et al. (2010)
HD 37776	91 ± 4	35 ± 4	1.5387115(9)	48857.124(3) ^p	sp	–	Mikulášek et al. (2008)
HD 44743	20 ± 7	41 ± 4	–	–	–	–	Fossati et al. (2015a), This work
HD 46328	≤8	15 ± 5	>30 (yr)	44296(304)	s, m	29.0, 26.0	Shultz et al. (2017)
HD 52089	21 ± 2	47 ± 2	–	–	–	–	Fossati et al. (2015a)
HD 55522*	70 ± 2	11 ± 9	2.7292(3)	53000.5(2) ^p	p, m	18.9, 11.8	Briquet et al. (2004), This work
HD 58260*	3 ± 2	16 ± 2	–	–	–	–	–
HD 61556	58 ± 3	19 ± 10	1.9087(6)	55198.05(5)	p, s, m	6.4, 13.8, 17.1	Shultz et al. (2015a)
HD 63425*	3 ± 3	19 ± 2	163(4)	55285(13)	m	9.4	This work
HD 64740*	135 ± 2	18 ± 11	1.330205(3)	43498.43(7)	m	14.3	Bohlender et al. (1987), This work
HD 66522*	3 ± 2	12 ± 2	909(75)	47365(76) ^p	p	7.3	This work
HD 66665*	8 ± 2	10 ± 1	24.5(1)	55274.4(8)	m	14.7	Petit et al. (2013), This work
HD 66765*	95 ± 3	12 ± 16	1.6079(5)	55907.9(2)	p, s, m	4.02, 11.8, 12.5	Alecian et al. (2014), This work
HD 67621*	21 ± 5	8 ± 5	3.593(1)	55904.7(1)	m	34.3	Alecian et al. (2014), This work
HD 96446	6 ± 2	15 ± 15	23.38(3)	55734(5) ^s	s, m	–, 28.4	Järvinen et al. (2017)
HD 105382	74 ± 5	7 ± 9	1.295(1)	47866.0(1) ^p	p	19.2	Briquet, Aerts & De Cat (2001)
HD 121743	74 ± 4	10 ± 8	1.130170(9)	56760.855	m	–	Briquet et al. (in preparation)
HD 122451*	70 ± 10	20 ± 20	2.885(1)	55706.6(4)	m	5.4	This work
HD 125823*	16 ± 2	10 ± 5	8.8169(1)	43733.6(35)	p, m	17.7, 32.4	Catalano & Leone (1996), This work
HD 127381	62 ± 4	15 ± 6	3.0194(2)	47620.5(6) ^p	p, m	4.9, 7.4	Henrichs et al. (2012)
HD 130807	28 ± 5	16 ± 6	2.9533(1)	55707.06(9)	p, m	4.4, 25.3	Buysschaert et al. (in preparation)
HD 136504A	35 ± 10	18 ± 6	–	–	–	–	–
HD 136504B	35 ± 5	5 ± 5	–	–	–	–	–
HD 142184	288 ± 6	5 ± 5	0.508276(1)	47913.694(1) ^s	p, s, m	7.6, 10.4, 4.4	Grunhut et al. (2012a)
HD 142990*	122 ± 2	1 ± 1	0.978832(2)	43563.00(5) ^p	p, m	15.5	Bychkov et al. (2005), This work
HD 149277*	8 ± 3	15 ± 2	25.380(7)	56119.5(3)	m	34.8	This work
HD 149438	7 ± 3	2 ± 2	41.033(2)	53482(2)	m	66.7	Donati et al. (2006)
HD 156324	50 ± 10	10 ± 6	1.5805(3)	56128.89(9) ^s	s, m	22.6, 14.4	Shultz et al. (in press)
HD 156424*	7 ± 2	14 ± 5	2.8721(6)	56126.3(3)	m	15.8	This work
HD 163472	62 ± 5	10 ± 5	3.6388330(9)	44807.71(6) ^s	u	18.8	Neiner et al. (2003b)
HD 164492	138 ± 10	34 ± 10	1.36986(7)	56770.53(7)	s, m	11.6, 68.9	Wade et al. (2017)
HD 175362*	34 ± 4	13 ± 8	3.67381(1)	43733.04(8)	p, m	34.8, 38.3	Bohlender et al. (1987), This work
HD 176582	103 ± 7	15 ± 15	1.581984(3)	54496.694(2) ^s	s, m	6.9, 20.7	Bohlender & Monin (2011)
HD 182180	306 ± 5	1 ± 1	0.5214404(6)	54940.83(5) ^s	p, s	57.2, 37.7	Oksala et al. (2010); Rivinius et al. (2010)
HD 184927	8 ± 2	8 ± 1	9.53102(7)	55706.517(5)	s, m	–, 15.8	Yakunin et al. (2015)
HD 186205*	4 ± 4	12 ± 3	37.21(6)	55640(3)	m	10.4	This work
HD 189775*	58 ± 2	13 ± 7	2.6071(3)	56262.01(8)	p, m	32.7	Petit et al. (2013), This work
HD 205021	37 ± 2	25 ± 5	12.000750(9)	52366.3(1)	u, m	37.5, 13.0	Henrichs et al. (2013)
HD 208057*	105 ± 5	33 ± 11	1.3678(4)	55028.2(1)	m	11.6	This work
ALS 3694*	48 ± 3	22 ± 12	1.678(3)	56819.5(2) ^s	s	5.4	This work

(1963), based on a 10 Å mm $^{-1}$ spectrogram. It is not clear how such a high $v \sin i$ was obtained. In any case, it is inconsistent with the value determined here; our results are however consistent with the measurement published by Kozlova, Glagolevskij & Klochkova (1975), 57 km s $^{-1}$.

4.2 Rotation periods

Amongst magnetic stars with radiative envelopes, magnetic measurements are in general modulated entirely by a star's rotation, since their fossil magnetic fields do not exhibit cyclic variability of

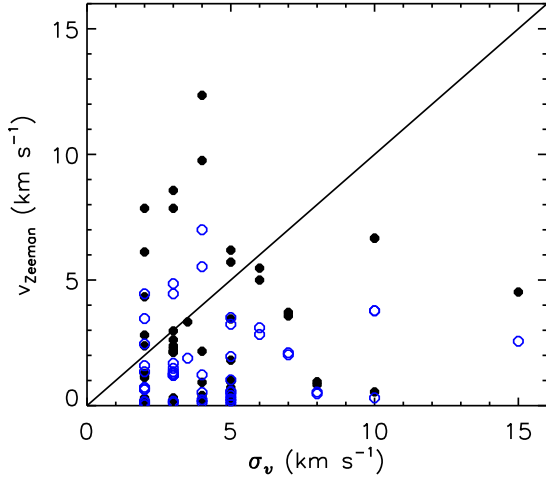


Figure 13. Zeeman broadening as a function of the uncertainty in $v \sin i$, σ_v . Filled black circles show the Zeeman broadening expected for a fictitious line with $\lambda_0 = 500$ nm and $g_{\text{eff}} = 1$. Open blue circles show the Zeeman broadening for the S II 566.5 nm line, with $g_{\text{eff}} = 0.5$.

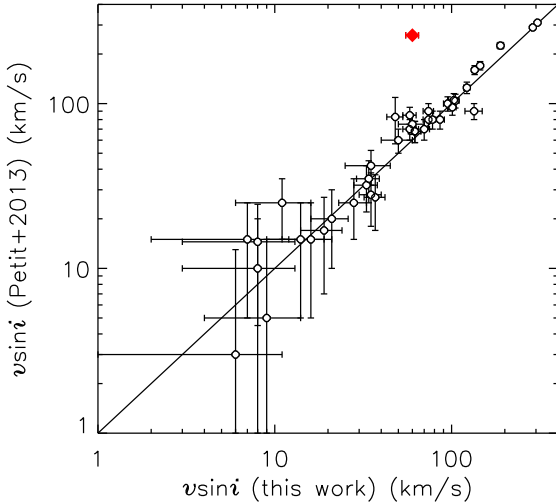


Figure 14. Comparison of $v \sin i$ measurements with literature values. Note that the axes are logarithmic. The solid red diamond indicates HD 35298, discussed further in the text.

the kind observed in cool stars with dynamo-generated surface magnetic fields. The same is true of their photometric and spectroscopic variabilities, since these arise in photospheric chemical spots and, in some cases, magnetospheric emission lines: in neither case has any variability been detected that could not be ascribed to rotation. The only exception to this rule of relevance to the magnetic early B-type stars are the β Cep and SPB pulsators, whose photometric and spectroscopic variability is dominated by pulsation; however, in these cases magnetic diagnostics are still primarily sensitive to rotation, e.g. Neiner et al. (2003a,b), Henrichs et al. (2013), and the time-scales of variability are also very different (hours for β Cep pulsation, days for rotation).⁶ Therefore, to determine rotation periods, in most cases we have relied on the $\langle B_z \rangle$ measurements

⁶ In the case of the β Cep pulsator ξ^1 CMa, a weak modulation of $\langle B_z \rangle$ with pulsation phase has been detected, however the amplitude of this modulation is much smaller than that of the rotational modulation, and this difference, combined with the very different time-scales, hours versus decades for this

described in Section 3. In two cases, ALS 3694 and HD 156324, we obtained periods using H α EWs, as both of these stars display magnetospheric H α emission (Alecian et al. 2014; Shultz et al. 2016), and the spectroscopic data sets are larger than the magnetic data sets. For HD 66765, we refined P_{rot} using *Hipparcos* photometry and He EWs, where we assumed both the photometric and the line profile variability in this He-strong star to be due to chemical spots. In one case, HD 66522, P_{rot} was determined using archival *Hipparcos* photometry, where once again we assume the origin of the photometric variability to be rotational modulation (the very long period determined for this star, ~ 900 d, is too long to be compatible with pulsation, while orbital modulation is unlikely as there is no evidence of RV variability). *Hipparcos* photometry was also used to refine P_{rot} in two further cases, HD 142990 and HD 35298.

Period analysis was performed using Lomb–Scargle statistics (Lomb 1976; Scargle 1982) as implemented in the IDL program PERIODOGRAM.PRO,⁷ which normalizes the periodogram to the total variance (Horne & Baliunas 1986). The uncertainty in each frequency was determined using the formula from Bloomfield (1976), $\sigma_F = \sqrt{6\sigma_{\text{obs}}/(\pi\sqrt{N_{\text{obs}}}A\Delta T)}$, where σ_{obs} is the mean uncertainty in the measurements, N_{obs} is the number of measurements, A is the amplitude of the RV curve, and ΔT is the timespan of observations.

In order to check that the rotation periods are physically plausible, period windows were bounded from above by

$$P_{\text{rot}} \leq \frac{2\pi R_{\text{eq}}}{v \sin i}, \quad (3)$$

and from below by the break-up velocity v_{br} (Jeans 1928), where R_{eq} is the equatorial radius. For the upper bound of the period window, the polar radius $R_p = R_{\text{eq}}$, while for the break-up velocity $R_{\text{eq}} = 1.5R_p$ due to rotationally induced oblateness. The majority of the stellar masses and radii used to determine v_{br} were obtained from Petit et al. (2013), with the exception of those stars in which magnetic fields were discovered subsequent to the publication of their catalogue. In these cases, masses and radii were obtained from Sikora et al. (2015) for HD 23478, Fossati et al. (2015a) for HD 44743 and HD 52089, Uytterhoeven et al. (2005) for HD 136504, and Wade et al. (2017) for HD 164492C.

Within a given period window, it is frequently the case that there are multiple peaks in the periodogram that phase the data more or less equally well, all of which meet the formal criterion for significance (an $S/N > 4$, Breger et al. 1993; Kuschnig et al. 1997). In order to check for spurious peaks, we used the $\langle N_z \rangle$ measurements discussed in Section 3.2. Peaks which appear in the $\langle N_z \rangle$ period spectrum are likely a consequence of the window function, and can be ignored. When $\langle N_z \rangle$ is not available (for historical $\langle B_z \rangle$ measurements as well as spectroscopic and photometric data), we used synthetic null measurements obtained via random Gaussian noise normalized to the mean uncertainty in order to derive the window contribution to the periodogram within the sampling window.

The statistical significance of a given peak in the periodogram can be quantified by means of the FAP, where we use equation (22) from Horne & Baliunas (1986) which gives the FAP as a function of the number of data points and the amplitude of the period spectrum. Similarly to the FAPs used in Section 3 to evaluate the statistical significance of the signal within Stokes V, smaller FAPs indicate

star, make their respective influences easily distinguishable (Shultz et al. 2017).

⁷ Available at <https://hesperia.gsfc.nasa.gov/ssw/gen/idl/util/periodogram.pro>

that a signal is less likely to be a consequence of white noise. We also calculated the S/N of each period, after pre-whitening with the most significant period (and the harmonics, for variations with higher order terms). No periods in Table 4 have an S/N below 4. The 13.6 d period for HD 44743 found by Fossati et al. (2015a) has an S/N of only 1.9, is thus likely to be spurious, and is therefore not included (and indeed, unpublished HARPSpol $\langle B_z \rangle$ measurements collected via the BRITepol LP are not coherently phased with this period; Buyschaert, private communication).

Depending on the size of a given data set, the time sampling, and the amplitude of the signal, in some cases there may be multiple peaks in the periodogram that meet the formal criterion for significance, and are of similar power, i.e. the rotational period given in Table 4 may not be unique. When this is the case, this is directly addressed in Appendix A, where for each star with a new or refined rotational period we show both periodograms, $\langle B_z \rangle$ curves, and where appropriate or available light curves phased with the adopted ephemeris. In general, we preferred solutions that implied higher rotational axis inclinations (i.e. longer periods), since higher inclinations are intrinsically more likely than lower inclinations.

In the end, new rotation periods have been determined for 10 stars. For a further 14 stars, comparison of the new magnetic data to $\langle B_z \rangle$ measurements in the literature has enabled refinement of the rotation periods. Rotation periods are given in Table 4. In general JD0 is taken to be the heliocentric Julian date at which $\langle B_z \rangle = |\langle B_z \rangle|_{\max}$ in the rotational cycle preceding the first observation in the time series, based upon a sinusoidal fit to the data. In some cases, maximum light or minimum EW were used to define JD0; in these cases, this is indicated with a superscript in the fifth column of Table 4. The uncertainty in JD0 was determined from the uncertainty in the phase of the sinusoidal fit. Also given in Table 4 are references for rotational periods obtained from the literature, and the method by which the period was obtained: via UV spectroscopy, optical spectroscopy, photometry, or magnetometry. When comparison of our measurements to historical data has allowed refinement of P_{rot} , this work is also given as a reference.

Rotation periods could not be determined for five stars: HD 44743, HD 52089, HD 58260, and HD 136504A and B. In the cases of HD 44743 and HD 136504, period analysis is hampered by both the small number of high-resolution $\langle B_z \rangle$ measurements, and by the low levels of variability relative to the median error bar in the data sets. For HD 52089, there are eight high-resolution $\langle B_z \rangle$ measurements available, in addition to *Hipparcos* photometry, but due to the very low level of variability in both data sets relative to the observational uncertainties no period above the S/N threshold of 4 could be identified. All of these stars have relatively weak $\langle B_z \rangle$ in comparison to the majority of the stars in the sample, around 10–100 G. In the case of HD 58260, despite the mean $\langle B_z \rangle$ of ~ 1.8 kG being significant at the 30σ level, the standard deviation of $\langle B_z \rangle$ is less than the mean 1σ error bar.

In two cases, HD 37479 and HD 37776, non-linear ephemerides are available that account for the observed spin-down of the star (Townsend et al. 2010; Mikulášek et al. 2008). Oksala et al. (2012) demonstrated the agreement achieved by this ephemeris between historical and modern $\langle B_z \rangle$ measurements of HD 37479. On their own, ESPaDOnS and Narval data are unable to distinguish between the spin-down ephemeris provided by Mikulášek et al. (2008) for HD 37776 and the newer ephemeris, in which spin-up of the star is reported (Mikulášek et al. 2011). We therefore adopt the earlier ephemeris as being the more conservative option, as shown in Fig. 11. This small ambiguity in ephemeris has no impact on the magnetic modelling.

The $\langle B_z \rangle$ curves of stars for which $\langle B_z \rangle$ curves have not already been published, or for which detailed individual studies are not currently in preparation, are shown in Appendix A. These are indicated with asterisks next to the star name in Table 4.

5 DISCUSSION

5.1 Selection of $\langle B_z \rangle$ data sets for modelling

The final column of Table 2 gives the type of measurement to be favoured for future modelling, based on the arguments below: ‘Z’ (LSD profiles extracted using line masks with metallic lines), ‘YZ’ (LSD profiles extracted using line masks with metallic and He lines), or ‘X’ (H lines).

With an ideal data set in which S/N is not a limitation, H line measurements would be used in all cases in order to avoid distortions due to chemical spots. However, for most stars the mean $\langle B_z \rangle$ significance Σ_B is lower in H lines than for LSD profiles, i.e. H line measurements are generally less precise than those obtained with LSD profiles. Furthermore, in the majority of cases differences in $\langle B_z \rangle$ measured from different elements are negligible (Fig. 9). H line measurements were thus selected only when $A_e/\sigma_e \geq 2$.

For stars with $A_e/\sigma_e \leq 2$, measurements obtained from LSD profiles extracted using metallic line masks are in general preferred, as these are unaffected by the extra broadening introduced by He lines. However, for stars with $\Sigma_B \sim 1$, $\langle B_z \rangle$ measurements using all available spectral lines were selected, as in these cases meaningful measurements are only possible when the maximum possible precision is achieved.

5.2 Fits to $\langle B_z \rangle$ curves

Several of the Stokes V profiles in Figs 2 and 3 show distortions that might be ascribed to non-dipolar surface magnetic fields, e.g. HD 35298, HD 36526, HD 61556, HD 127381, HD 142990, HD 156324, HD 176582, and HD 189775. However, such distortions might also easily arise from chemical spots or pulsations. As a pure centred dipole should produce a sinusoidal $\langle B_z \rangle$ variation, in order to determine the dominant magnetic topologies of the sample, we have modelled the $\langle B_z \rangle$ curves by fitting least-squares sinusoids, using the equation

$$\langle B_z \rangle(\phi) = \sum_{n=0}^3 B_n \sin(n\phi + \Phi_n), \quad (4)$$

where ϕ is the rotational phase determined using the ephemerides in Table 4, B_n is the amplitude of the n th-order sinusoid, and Φ_n is the phase offset. We utilized the IDL routine CURVEFIT, which computes the least-squares solution and its 1σ uncertainties to an arbitrary non-linear function so long as its partial derivatives are known. We kept B_n and Φ_n as free parameters, initially setting B_0 to the mean $\langle B_z \rangle$, B_1 – B_3 as the standard deviation of $\langle B_z \rangle$ divided by the order, and $\Phi_n = 0$. This simple model is meant to reproduce approximately the expected $\langle B_z \rangle$ behaviour of dipolar, quadrupolar, and octupolar magnetic field components. The coefficients B_n are given in Table 5. Including a phase offset between the components accounts for possibility that they do not all share the same tilt angle with respect to the rotational axis.

We utilized the same $\langle B_z \rangle$ measurements chosen for modelling in Section 5.1 (see also Table 2). In the case of HD 105382, there are too few high-resolution $\langle B_z \rangle$ measurements (only three) to constrain even a first-order fit, therefore for this star, the fit was calculated

Table 5. Curve fitting of $\langle B_z \rangle$ curves. Stars for which FORS data were used to constrain the fit are indicated with an asterisk next to the HD number in the first column. The second column gives the significance Σ_{Amp} of the $\langle B_z \rangle$ amplitude, as explained in the text. The third column gives the reduced χ^2 of a single-order sinusoidal fit. The fourth through seventh columns gives the amplitudes B_n of the n th-order least-squares sinusoidal fits.

HD number	Σ_{Amp}	χ^2/ν	B_0 (kG)	B_1 (kG)	B_2 (kG)	B_3 (kG)
3360	6.5	1.0	-0.0037 ± 0.0006	0.0098 ± 0.0009	—	—
23478	4.3	1.5	-0.98 ± 0.03	0.14 ± 0.05	—	—
25558	3.4	0.5	-0.045 ± 0.006	0.044 ± 0.008	—	—
35298	43	4.8	0.40 ± 0.04	3.40 ± 0.06	0.17 ± 0.05	0.22 ± 0.06
35502	17	0.6	-1.22 ± 0.03	1.66 ± 0.04	—	—
36485	3.9	1.4	-2.48 ± 0.02	0.06 ± 0.03	0.08 ± 0.02	—
36526	27	1.3	1.34 ± 0.05	2.25 ± 0.08	—	—
36982	3.8	0.5	0.129 ± 0.008	0.09 ± 0.01	—	—
37017	12	0.9	-0.80 ± 0.03	0.98 ± 0.05	—	—
37058	33	1.5	-0.19 ± 0.02	0.77 ± 0.03	—	—
37061	9.6	1.7	-0.11 ± 0.02	0.17 ± 0.02	—	—
37479	37	19	0.67 ± 0.05	2.37 ± 0.07	0.68 ± 0.07	—
37776	18	39	-0.19 ± 0.04	0.35 ± 0.05	0.96 ± 0.06	0.96 ± 0.06
46328	99	1.1	-0.21 ± 0.01	0.54 ± 0.01	—	—
55522	13	1.0	-0.01 ± 0.04	0.83 ± 0.06	—	—
61556	33	8.3	-0.37 ± 0.01	0.45 ± 0.01	0.07 ± 0.02	0.06 ± 0.02
63425	11	1.1	0.095 ± 0.002	0.050 ± 0.003	—	—
64740	27	3.0	-0.21 ± 0.01	0.57 ± 0.02	0.07 ± 0.02	—
66522	23	0.9	-0.02 ± 0.02	0.74 ± 0.01	—	—
66665	14	0.7	-0.040 ± 0.002	0.081 ± 0.002	—	—
66765	18	1.3	-0.17 ± 0.01	0.56 ± 0.02	—	—
67621	13	1.0	0.160 ± 0.004	0.250 ± 0.006	—	—
96446	20	0.9	-0.74 ± 0.04	0.31 ± 0.02	—	—
105382*	8.5	4.5	-0.33 ± 0.04	0.34 ± 0.06	—	—
121743	3.1	1.0	0.26 ± 0.01	0.06 ± 0.03	—	—
122451	9.3	3.0	-0.024 ± 0.002	0.010 ± 0.004	—	—
125823	18	1.2	0.03 ± 0.01	0.48 ± 0.02	—	—
127381	7.6	1.4	0.016 ± 0.009	0.12 ± 0.01	—	—
130807	48	2.6	0.384 ± 0.006	0.626 ± 0.008	0.050 ± 0.008	—
142184	11	3.6	-1.55 ± 0.03	0.54 ± 0.05	0.24 ± 0.04	—
142990	20	3.2	-0.12 ± 0.03	1.16 ± 0.04	0.20 ± 0.05	—
149277	24	1.2	0.41 ± 0.05	2.35 ± 0.07	—	—
149438	34	38	0.0094 ± 0.0009	0.049 ± 0.001	0.032 ± 0.001	0.019 ± 0.001
156324	10	1.4	0.90 ± 0.07	1.18 ± 0.09	—	—
156424	6.7	1.4	-0.34 ± 0.02	0.17 ± 0.03	—	—
163472	8.7	1.2	-0.029 ± 0.004	0.080 ± 0.006	—	—
164492	13	0.8	0.75 ± 0.03	1.07 ± 0.04	—	—
175362	155	307	-1.02 ± 0.02	4.89 ± 0.03	0.93 ± 0.03	0.22 ± 0.03
176582	28	3.2	-0.00 ± 0.02	1.49 ± 0.03	0.05 ± 0.03	0.15 ± 0.03
182180	44	1.8	0.24 ± 0.08	2.44 ± 0.09	—	—
184927	37	1.3	0.92 ± 0.01	0.80 ± 0.02	—	—
186205	9.3	1.1	-0.717 ± 0.009	0.12 ± 0.02	—	—
189775	25	14	0.38 ± 0.02	0.61 ± 0.02	0.23 ± 0.02	—
205021	19	0.4	0.003 ± 0.002	0.102 ± 0.003	—	—
208057	8.0	0.8	-0.001 ± 0.005	0.090 ± 0.008	—	—
ALS 3694*	1.8	0.8	-1.85 ± 0.05	0.2 ± 0.1	—	—

together with the FORS1 $\langle B_z \rangle$ measurements performed by Bagnulo et al. (2015). The FORS1 $\langle B_z \rangle$ measurements published by Bagnulo et al. (2006) were also used to help constrain the fit for ALS 3694, due to the very large error bars of the ESPaDOnS $\langle B_z \rangle$ measurements relative to the variation in $\langle B_z \rangle$.

The most straightforward way to look for the presence of non-dipolar surface magnetic fields is to evaluate the goodness-of-fit χ^2/ν of the best-fitting first-order sinusoidal curves, where ν is the number of degrees of freedom. If $\chi^2/\nu \sim 1$, the model is a reasonable fit to the data, indicating that $\langle B_z \rangle$ is consistent with a

centred dipole. Examples are shown in Fig. 15 for HD 36526, for which a first-order sinusoid is a good fit, and HD 189775, which is not fit well by this model. For those stars for which $\chi^2/\nu \sim 1$, only B_0 and B_1 are given in Table 5, since fitting higher order terms would tend to reduce χ^2/ν below 1 by overfitting to the noise in the data. Since the reduced χ^2 is sensitive to the precision of the data set, we compared the goodness of fit to the significance of the amplitude of the $\langle B_z \rangle$ curve, $\Sigma_{\text{Amp}} = |\langle B_z \rangle_{\text{max}} - \langle B_z \rangle_{\text{min}}|/\sigma_B$, where σ_B is the median $\langle B_z \rangle$ error bar. It should be recalled that the mean significance of $\langle N_z \rangle$ $\Sigma_N < 1$ for all stars, indicating that the $\langle B_z \rangle$

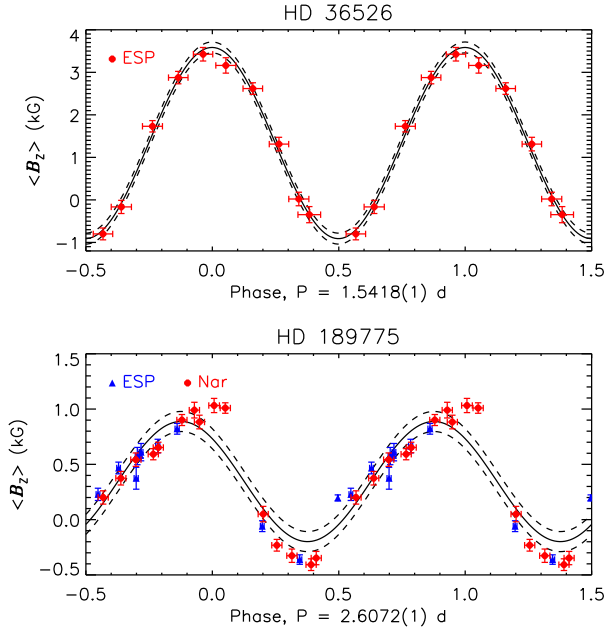


Figure 15. First-order sinusoidal fits for HD 36526 (which yields $\chi^2/\nu \sim 1$) and HD 189775 (with $\chi^2/\nu \sim 20$). Solid curves show the fits; dashed curves show the 1σ uncertainty in the fits. The legend indicates the source of the data, ESP(aDONs), or Nar(val).

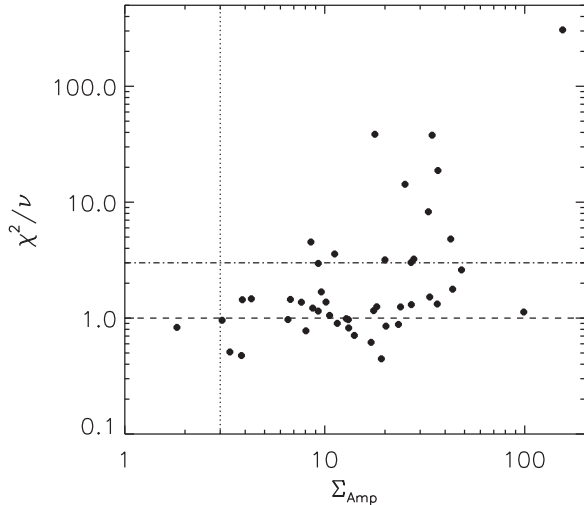


Figure 16. Goodness-of-fit χ^2/ν of first-order fits to $\langle B_z \rangle$ as a function of the significance of the $\langle B_z \rangle$ variation Σ_{Amp} . Points to the right of the dotted line possess a $\langle B_z \rangle$ variation of at least 3σ significance. The dashed line indicates $\chi^2/\nu = 1$, the formal definition of a ‘good’ fit. The dot-dashed line indicates $\chi^2/\nu = 3$. Points above the dot-dashed line and to the right of the dotted line show the strongest evidence for higher order multipole contributions to $\langle B_z \rangle$.

error bars are consistent with the scatter in the data, hence values of $\chi^2/\nu > 1$ should reflect real departures from sinusoidal variations. χ^2/ν and Σ_{Amp} are given in Table 5.

Fig. 16 shows χ^2/ν as a function of Σ_{Amp} . The comparison is performed for 46/51 stars, with the remainder left out due to the absence of a firmly established P_{rot} . There is an approximate correlation between the two measures, suggesting that, as expected, departures from purely sinusoidal behaviour are more easily detected in data sets with a higher precision relative to the strength

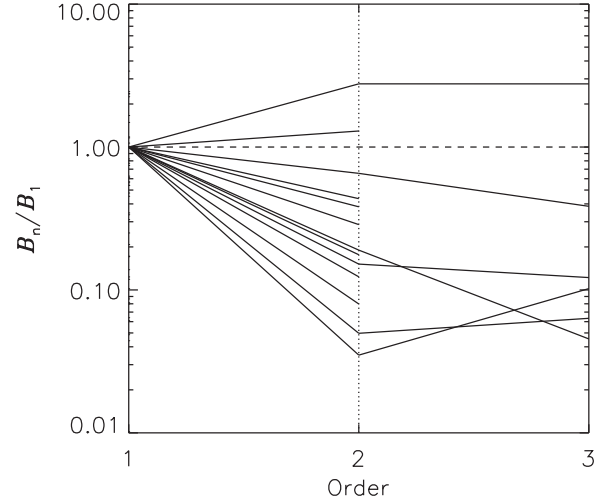


Figure 17. Coefficients B_n normalized to B_1 , for the stars for which first-order sinusoids yield $\chi^2/\nu > 3$.

of the magnetic field. Only one star, ALS 3694, has $\Sigma_{\text{Amp}} \leq 3$: for this star, a departure from sinusoidal behaviour cannot be evaluated, since the quality of the data is too low to reliably constrain a higher order fit. Of the 45 stars with $\Sigma_{\text{Amp}} \geq 3$, 33 have $\chi^2/\nu \sim 1$, indicating that a single-order sinusoid provides a reasonably good fit to the data.

The remaining 12 stars have $\chi^2/\nu \geq 3$, suggesting a dipolar model may not be appropriate in these cases. In all cases, H line measurements were used to evaluate χ^2/ν , therefore it is unlikely that this result is merely a consequence of distortion of $\langle B_z \rangle$ by chemical spots. Of these, six stars are already well known to exhibit significant departures from the sinusoidal behaviour expected of a centred dipole: HD 175362 (Wolff & Wolff 1976; Borra et al. 1983), HD 37776 (Thompson & Landstreet 1985), HD 37479 (Townsend, Owocki & Groote 2005), HD 149438 (Donati et al. 2006), HD 142184 (Grunhut et al. 2012a), and HD 61556 (Shultz et al. 2015a). Indeed, HD 37776, HD 37479, and HD 149438 have been verified to possess multipolar magnetic fields via ZDI (Kochukhov et al. 2011; Oksala et al. 2015b; Donati et al. 2006). The remainder, in increasing order of χ^2/ν , are: HD 64740, HD 142990, HD 176582, HD 105382, HD 35298, and HD 189775. Since the fit to HD 105382 was constrained using FORS1 data, the high value of χ^2/ν likely reflects systematic differences between FORS1 and HARPSpol, rather than intrinsic magnetic complexity (Landstreet, Bagnulo & Fossati 2014).

Examination of χ^2/ν can identify poor fits, but it does not necessarily indicate that the departure from purely sinusoidal behaviour is either large or systematic. In Fig. 17, the coefficients B_n are shown as a function of n , normalized to B_1 . In all but three cases $B_2 \leq 0.5B_1$, and it is often much smaller than this. In only two cases is $B_2 > B_1$: HD 37776, and HD 36485 (which, although it is well fitted by a first-order sinusoid, we fitted with a second-order sinusoid for reasons explained below). B_3 is typically less than or similar to B_2 , and is larger than B_1 only for HD 37776, for which star ZDI has found the magnetic topology to be dominated by the $\ell = 2$ mode (Kochukhov et al. 2011). This suggests that even for the majority of the stars with $\langle B_z \rangle$ curves suggestive of some degree of magnetic complexity, their surface magnetic fields are still dominated by the dipolar component.

The $\langle B_z \rangle$ curve of one star, HD 36485, yields relatively low values of the above measures ($\Sigma_{\text{Amp}} = 3.9$ and $\chi^2/\nu = 1.4$). While the

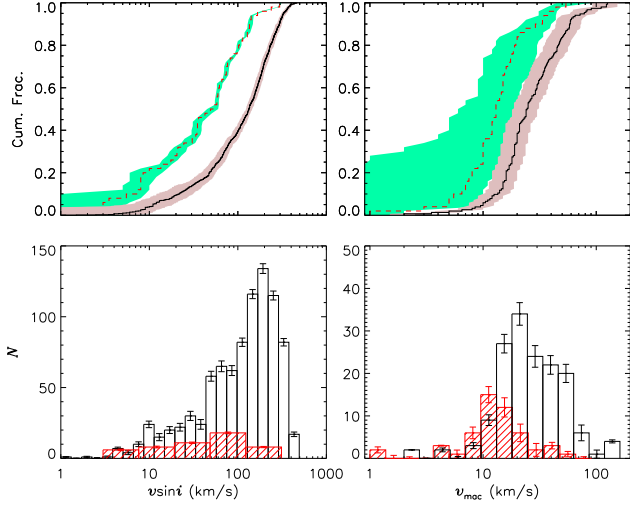


Figure 18. Cumulative distributions (top) and histograms (bottom) of $v \sin i$ (left) and v_{mac} (right) for the magnetic B-type stars in this study (dashed and hatched red) and main-sequence B-type stars from the literature (black). Shaded regions in the cumulative distributions indicate uncertainties.

low level of variation, likely a consequence of its rotational pole being nearly aligned with the line of sight (Leone et al. 2010), means that the signature of a complex surface magnetic field is not formally significant according to the criteria adopted here, close examination of its $\langle B_z \rangle$ curve suggests a double-wave variation that would indicate significant contribution from multipolar components (see Fig. A3). This is the only star other than HD 37776 and HD 149438 whose $\langle B_z \rangle$ curve shows more than two extrema.

We conclude that, out of the 46 stars in the sample with rotation periods, we are not able to detect any evidence in $\langle B_z \rangle$ of magnetic field topologies more complex than a centred dipole in 34 of them. This is despite the quality of the data being sufficient to do so for 33 of the stars with $\langle B_z \rangle$ curves consistent with a dipolar magnetic field. Amongst the 12 stars showing evidence of some degree of magnetic complexity, in all but three cases the dipolar component is still dominant. Thus, the majority of the sample are well described by dipolar models. Therefore, non-dipole components of the magnetic field of sufficient magnitude to strongly distort $\langle B_z \rangle$ are relatively uncommon, and even when detectable, generally modify $\langle B_z \rangle$ to only a small degree. However, we have detected evidence in the $\langle B_z \rangle$ curves of multipolar magnetic field components in six stars for which this has not previously been reported: HD 35298, HD 36485, HD 64740, HD 142990, HD 176582, and HD 189775.

5.3 Rotational and turbulent broadening in magnetic versus non-magnetic stars

Fig. 18 compares the $v \sin i$ and v_{mac} measurements of the magnetic B-type stars in the present sample to those reported for non-magnetic, main-sequence B-type stars in the literature, where for the literature stars $v \sin i$ measurements were taken from Dufton et al. (2013), Simón-Díaz & Herrero (2014), Garmy et al. (2015), and Simón-Díaz et al. (2017) (a total of 890 $v \sin i$ measurements of non-magnetic B-type stars), while v_{mac} measurements were provided only by Simón-Díaz & Herrero (2014) and Simón-Díaz et al. (2017) (for a total of 154 v_{mac} measurements of non-magnetic B-type stars). Bin sizes were determined using the Freedman–Diaconis rule, which optimizes the bin size for the variance and size of the data set (Freedman & Diaconis 1981). Error bars arise from the

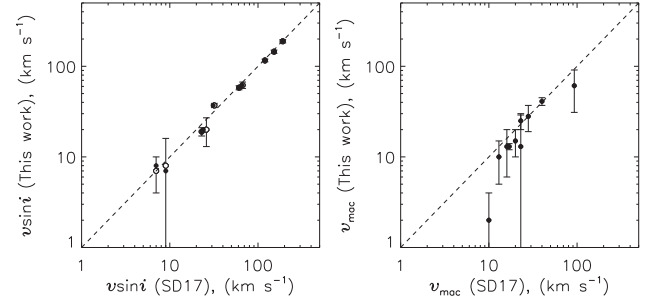


Figure 19. Comparison of $v \sin i$ (left) and v_{mac} (right) for the stars appearing in both the present sample and in that of Simón-Díaz et al. (2017). Open circles indicate Fourier transform measurements, filled circles indicate goodness-of-fit measurements.

1 σ uncertainties in the individual measurements. 10 000 simulated data sets were created, of the same size as the original data set, with values of individual data points equal to the original measurement plus a random number generated from a Gaussian distribution normalized to the 1 σ uncertainty in the original measurement. Histograms were then generated from each simulated data set, and the error bar in each histogram bin was obtained from the standard deviation across the histograms for all synthetic data sets. Cumulative distribution uncertainties were obtained from the cumulative distributions of the measurements plus or minus their 1 σ error bars.

The literature $v \sin i$ measurements combine extragalactic stars (the VLT–FLAMES Tarantula Survey, Dufton et al. 2013) and Galactic stars (Simón-Díaz & Herrero 2014; Garmy et al. 2015; Simón-Díaz et al. 2017). The samples presented by Simón-Díaz & Herrero (2014) and Simón-Díaz et al. (2017) overlap slightly with one another and with the present sample; stars appearing in both were removed from the Simón-Díaz & Herrero (2014) and Simón-Díaz et al. (2017) data, and for stars appearing in both Simón-Díaz & Herrero (2014) and Simón-Díaz et al. (2017) the more recent values were used. Fig. 19 compares the $v \sin i$ and v_{mac} values determined here to those determined for the same stars by Simón-Díaz et al. (2017). There is good agreement for $v \sin i$, with no appearance of systematic differences between the two samples. However, our values of v_{mac} are systematically lower by about 20 per cent. This is likely due to the use of a δ -function for the intrinsic profile by Simón-Díaz & Herrero (2014) and Simón-Díaz et al. (2017), whereas we elected to use Gaussian intrinsic profiles with FWHM values determined by the thermal broadening of the line, typically around a few km s^{-1} . Experiments for a few stars in which δ -functions were utilized instead recovered values of v_{mac} comparable with those found by Simón-Díaz & Herrero (2014) and Simón-Díaz et al. (2017).

The two other samples presumably contain some number of magnetic stars, but the visual magnitudes of these stars are too faint for high-resolution magnetometry to distinguish magnetic from non-magnetic stars; in any case, the fraction of magnetic stars should be no higher than 10 per cent. Comparing $v \sin i$ values, there is a clear difference between magnetic and non-magnetic B-type stars, with the distribution of the former peaking at under 100 km s^{-1} , while the latter peaks at about 200 km s^{-1} . A two-sample Kolmogorov–Smirnov (K-S) test yields a probability of 7×10^{-8} that the two samples belong to the same distribution. This is in line with expectations that magnetic stars should be systematically more slowly rotating than non-magnetic stars due to braking by the magnetized stellar wind. It should be noted that this difference may be even more pronounced than suggested by Fig. 18: Dufton et al. (2013)

noted the distribution of $v \sin i$ in the Large Magellanic Cloud B-type stars is bimodal, with a broad-lined component (~ 75 per cent of the sample) with $v \sin i \sim 250 \text{ km s}^{-1}$, and a narrow-lined component with $v \sin i < 100 \text{ km s}^{-1}$. Dufton et al. speculated that the narrow-lined population is their sample might be a consequence of magnetic braking, noting that the fraction expected to be spun down by tidal interactions with a binary companion (the other leading mechanism for producing slow rotation) is much smaller than the observed fraction of slow rotators (de Mink et al. 2013). Nevertheless, given that Dufton et al. (2013) had no evidence beyond rotation to infer magnetic fields in the slowly rotating sub-sample, and that removing these stars from the comparison would artificially increase the difference between stars with and without confirmed magnetic fields, we have made the conservative assumption that the sub-sample of narrow-lined stars are not magnetic, and so have included them in the comparison. It is worth noting that the velocity peak of their narrow-lined component is similar to that found here for magnetic stars, suggesting that these stars may be worthy of follow-up observations to determine if they exhibit the spectroscopic peculiarities and periodic photometric modulations expected for magnetic stars. However, both MiMeS and BOB have observed a large number of sharp-lined Galactic stars, many of which are not detectably magnetic (Fossati et al. 2015b; Wade et al. 2016; Grunhut et al. 2017).

Comparison of v_{mac} values determined here to those presented for non-magnetic B-type stars by Simón-Díaz & Herrero (2014) (right-hand panels of Fig. 18) yields a somewhat more ambiguous result. The probability that the two samples belong to the same distribution according to the K-S test is 10^{-10} , a low enough value to suggest the difference is real. The magnetic sample peaks at a lower value ($\sim 10 \text{ km s}^{-1}$) than the non-magnetic sample ($\sim 20 \text{ km s}^{-1}$), with the non-magnetic stars having generally higher values of v_{mac} : no magnetic star has $v_{\text{mac}} > 60 \text{ km s}^{-1}$, while the non-magnetic sample extends up to 100 km s^{-1} ; conversely, the lowest value in the non-magnetic sample is around 10 km s^{-1} , while several of the magnetic stars have v_{mac} consistent with 0 km s^{-1} . It should be noted that in the sample of sharp-lined early B-type stars studied by Nieva & Przybilla (2012), several stars with $v_{\text{mac}} = 0 \text{ km s}^{-1}$ were found. Only one of the stars (β Cep) studied by Nieva & Przybilla (2012) is known to possess a magnetic field, and the value of v_{mac} they found ($20 \pm 7 \text{ km s}^{-1}$) is compatible with the value found here ($25 \pm 5 \text{ km s}^{-1}$). However, the Nieva & Przybilla (2012) sample was selected for stars with sharp spectral lines, is thus intrinsically biased towards small line-broadening parameters, and therefore may not give an accurate representation of the true distribution of macroturbulent velocities in main-sequence early B-type stars.

While the physical origin of macroturbulence is still debated, the leading possibility is that it arises from the superposition of multiple high-order non-radial g -mode pulsations excited in the sub-surface Fe convection zone (Aerts et al. 2009; Godart et al. 2017). Strong magnetic fields are expected to suppress convection and, hence, inhibit macroturbulence, an effect which has been observed in the strongly magnetized Of?p star NGC 1624-2 (Sundqvist et al. 2013a). If the magnetic field plays a role in suppressing macroturbulence in the present sample, we should expect to see that more strongly magnetized stars tend to have lower values of v_{mac} . This comparison is performed in Fig. 20. A least-squares linear fit finds a trend of decreasing v_{mac} with increasing $|\langle B_z \rangle|_{\text{max}}$, which is robust against iterative removal of individual data points as indicated by the shaded region. The Pearson's correlation coefficient for v_{mac} versus $\log |\langle B_z \rangle|_{\text{max}}$ is -0.36 ± 0.03 .

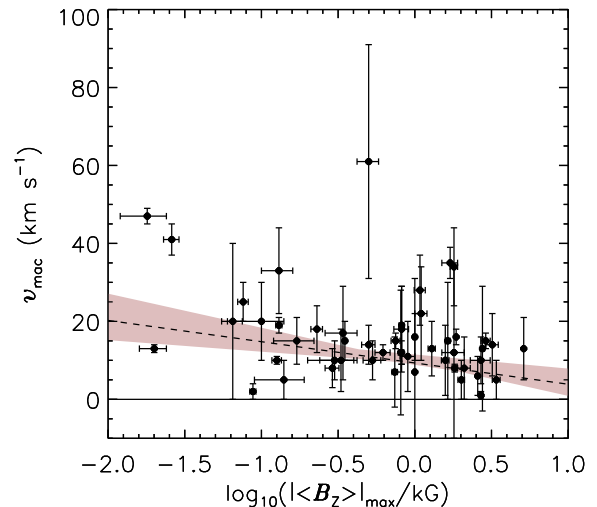


Figure 20. v_{mac} as a function of $\log(|\langle B_z \rangle|_{\text{max}})$. The dashed line indicates the least-squares linear fit, and the shaded region shows the uncertainty in the fit derived from iterative removal of individual data points.

The much smaller difference in v_{mac} distributions as compared to $v \sin i$ could also be a consequence of subtle differences in the line-profile fitting routines used here and employed by Simón-Díaz & Herrero. There is evidence for such a systematic discrepancy in Fig. 19, which could explain part or all of the difference. The median ratio of v_{mac} values from Simón-Díaz & Herrero to those found here (Fig. 19) is 1.3, while the median ratio between the cumulative distributions in Fig. 18 is 2.2, suggesting that the different choices of intrinsic profiles may not explain all of the differences in the distributions. Uncertainties in v_{mac} tend to be higher than those in $v \sin i$ (around 25 per cent as compared to around 7 per cent), and as can be seen from the uncertainties in the cumulative distributions of the magnetic and non-magnetic samples, they overlap at around 1σ . The results of Nieva & Przybilla (2012), who found $v_{\text{mac}} = 0 \text{ km s}^{-1}$ for several early B-type stars not known to possess a magnetic field, further suggests that methodological differences (non-LTE spectral fitting versus line-profile fitting and Fourier transforms) may underlie the different results obtained in their study, versus that of Simón-Díaz & Herrero. Verifying whether or not this difference in macroturbulent velocities is real will require a self-consistent analysis of the combined sample of magnetic and non-magnetic main sequence B-type stars. It may also benefit from a more sophisticated treatment of the line profile variability of the chemically peculiar stars in this sample. In particular, accounting for the effect of chemical spots on line profile shapes may be necessary, as it is possible these distortions may be misinterpreted as macroturbulence by the method used here [as was found by Aerts et al. (2014) for HD 105382, for which v_{mac} was a strong function of rotational phase]. If this is the case, the values of v_{mac} in Table 4 will tend to be higher than the true v_{mac} , thus accentuating the difference between the magnetic and non-magnetic stars. Such a comparison should also utilize predictions based upon stellar structure models of the depth to which the surface magnetic fields are expected to inhibit convection in any given star.

5.4 Comparison of rotational periods in magnetic early B-type stars to Ap/Bp stars and magnetic O-type stars

Fig. 21 compares the cumulative distribution of rotation periods for the early B-type stars, to the cumulative distribution of the rotational

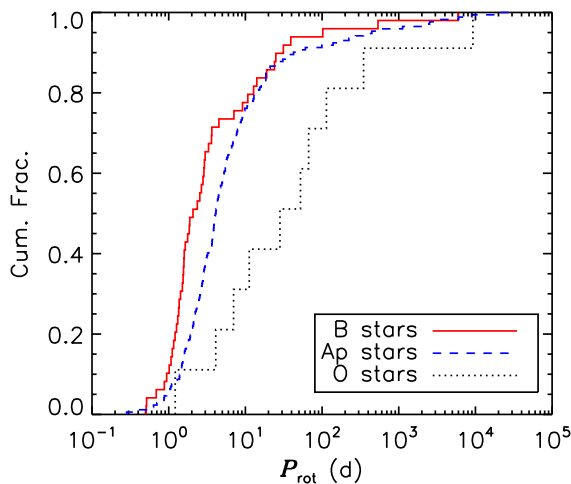


Figure 21. Cumulative distributions of the rotational periods of magnetic early B-type stars, Ap and late-type Bp stars, and magnetic O-type stars.

periods for Ap stars and late-type Bp stars, where the Ap/Bp stars have been compiled from the catalogues published by Landstreet & Mathys (2000), Bychkov et al. (2005), Aurière et al. (2007), and Power (2007), for a total comparison sample of 172 stars. As there is some overlap between the samples, particularly with the comprehensive Bychkov et al. catalogue, each sample was cleaned by HD number to ensure no duplication. The overall distributions are quite similar: the median rotation period of both the present sample and the cooler stars is on the order of days, and both possess tails extending to around 10^4 d. However, there appears to be a higher percentage of rapidly rotating stars amongst the early-type stars.

To evaluate whether or not the difference between the two populations is significant, we performed a K-S test. The K-S test significance is 0.0017, which is quite small and suggests the two distributions may indeed be different. This is a somewhat unexpected result, as given the stronger stellar winds and consequently more rapid magnetic braking experienced by early-type stars, they should presumably be, in general, slower rotators than the late-type stars. The difference may be indicative that the present sample is generally younger than the late-type stars, and/or that they were born with substantially more angular momentum than the Ap/Bp stars. Alternatively, this could be an observational bias arising from the inclusion in this sample of stars that were first identified as spectropolarimetric targets based upon emission signatures suggestive of a centrifugal magnetosphere (Townsend & Owocki 2005; Petit et al. 2013), e.g. HD 23478 (Sikora et al. 2015; Hubrig et al. 2015).

Fig. 21 also compares the distribution of rotational periods in the magnetic B-type stars to those of the magnetic O-type stars, which we have obtained from those collected by Petit et al. (2013), with the exception of HD 47129 which was provided by Grunhut (private communication). While the extrema of the O-type star period distribution are similar to those of the other two distributions, the magnetic O-type stars are systematically more slowly rotating than either the Ap/Bp stars or the early B-type stars. The K-S test yields a significance of 0.0011 that the B- and O-type stars belong to different distributions. The slower rotation of the magnetic O-type stars is an expected result since, due to their powerful stellar winds, they should lose angular momentum much more rapidly than cooler stars.

6 CONCLUSIONS AND FUTURE WORK

We measured $\langle B_z \rangle$ using both LSD profiles and H lines. Comparison of $\langle B_z \rangle$ measurements obtained from H lines and LSD profiles extracted using single-element masks has revealed that many of the Bp stars display systematic element-dependent differences in the shapes and amplitudes of their $\langle B_z \rangle$ curves, confirming this phenomenon as essentially ubiquitous in this class of stars. Chemically normal stars, while a minority of the sample, show no evidence of this effect. $\langle B_z \rangle$ measurements obtained using H lines show excellent agreement with historical data, further confirming the reliability of ESPaDOnS, Narval, and HARPSpol magnetometry, and in several cases enabling more precise rotational periods to be determined via comparison of the new $\langle B_z \rangle$ curves to older measurements.

No magnetic field was detected in HD 35912, indicating that the star has been incorrectly identified as magnetic in the literature, and reducing the size of the sample from 52 to 51 stars. Using magnetic, spectroscopic, and archival photometric data, we have determined new rotational periods for 10 of 51 stars. No period could be determined for five stars: HD 44743, HD 52089, HD 58260, and the two components of the doubly magnetic close binary HD 136504. The remarkable stability of the $\langle B_z \rangle$ measurements of HD 58260 over the 35 yr of observations suggests either that the star's rotational period is extremely long, and/or that its rotational pole is almost perfectly aligned with the line of sight, and/or that its magnetic field is axisymmetric with the magnetic axis almost perfectly aligned with the rotational axis. For the remaining stars without rotational periods, the available data sets are simply too small and/or imprecise to identify a period, but all show some degree of variability and there is no reason to think that their rotational periods cannot be determined if additional, higher S/N data are obtained.

The distributions of rotational periods appear to be broadly similar between magnetic early B-type stars, and Ap/late-type Bp stars. There is some suggestion that the hotter stars possess an excess of rapid rotators compared to the cooler stars, a discrepancy that could be reflective either of observational bias or a real difference in the populations arising from e.g. differing angular momentum distributions at birth.

Projected rotational broadening $v \sin i$ was measured together with a radial/tangential turbulent broadening parameter v_{mac} , with multiple spectral lines used for each star in order to obtain statistical uncertainties in the line broadening parameters. Zeeman splitting is able to account for the majority of the line broadening in HD 58260 and HD 96446, both of which have strong magnetic fields and sharp spectral lines. Comparison to previous $v \sin i$ measurements revealed only one 3σ outlier, HD 35298, for which we revise $v \sin i$ sharply downwards. Comparing the distribution of $v \sin i$ to non-magnetic stars shows that $v \sin i$ peaks about 100 km s^{-1} below the peak for non-magnetic stars, consistent with the magnetic stars having lost angular momentum more quickly due to magnetic braking. A similar comparison of v_{mac} suggests that turbulent broadening may also be systematically lower in magnetic stars, although this conclusion is only tentative and should be investigated with a larger sample measured with a self-consistent methodology. We find a trend of decreasing v_{mac} with increasing $|\langle B_z \rangle_{\text{max}}|$, which tends to support the conclusion that strong magnetic fields suppress turbulence. If this proves to be the case, this may give some insight into the origin of the macroturbulence phenomenon in B-type stars in general.

We have searched for the signature of multipolar magnetic fields via comparison of $\langle B_z \rangle$ curves to the first-order sinusoidal variation

expected for a centred, tilted dipole. The $\langle B_z \rangle$ curves of 74 per cent of the sample (34 of 46 stars with rotational periods) are consistent with dipolar magnetic fields. Of the remaining 12 stars (26 per cent of the sample), the contributions to $\langle B_z \rangle$ from higher order multipoles are typically minor, with amplitude coefficients generally less than 20 per cent of the magnitude of the dipolar component. In addition to the six stars previously identified as having multipolar surface fields, we have found evidence for magnetic complexity in six stars for which it was not previously reported: HD 35298, HD 36485, HD 64740, HD 142990, and 176582, and HD 189775. While the reduced χ^2 of a first-order fit to HD 36485's $\langle B_z \rangle$ curve is formally consistent with a centred dipole, there is some evidence of a quadrupolar component, which requires verification at higher S/N given the very small level of variability exhibited by $\langle B_z \rangle$.

This is Paper I in a series that will explore the magnetic and magnetospheric properties of the magnetic early B-type stars. In Paper II, we will use the rotational periods and $\langle B_z \rangle$ measurements presented here to obtain dipolar oblique rotator model parameters together with magnetospheric parameters, with which we will revisit the rotation-magnetic confinement diagram introduced by Petit et al. (2013) with the aim of exploring (1) the boundary between stars with H α in emission versus absorption, and (2) the rotational evolution of these stars.

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APPENDIX A: ROTATION PERIODS AND LONGITUDINAL MAGNETIC FIELD MEASUREMENTS OF INDIVIDUAL STARS

In this appendix, we present $\langle B_z \rangle$ curves for stars with existing periods that we have refined using new data (14 stars), the stars for which we have determined new rotational periods using magnetic, spectroscopic, or photometric data (10 stars), one star for which we could not uniquely determine P_{rot} (HD 58260), and one star for which we have new $\langle B_z \rangle$ measurements but could not improve P_{rot} (HD 36485). We also present periodograms and, in some cases, photometric or spectroscopic data together with $\langle B_z \rangle$. The abscissae of the periodograms are restricted to the period windows discussed in Section 4.2. For stars with newly determined rotation periods, we state in the text whether we consider the period to be confidently established, or either tentative or only partially constrained.

Stars for which the existing magnetic data have already been published, or for which detailed individual studies are in preparation, are not included in this appendix. These are HD 3360 (Briquet et al. 2016), HD 23478 (Sikora et al. 2015), HD 35502 (Sikora et al. 2016), HD 37479 (Oksala et al. 2012), HD 44743 and HD 52089 (Fossati et al. 2015a; Neiner et al. 2017a), HD 46328 (Shultz et al. 2017), HD 61556 (Shultz et al. 2015a), HD 105382 (Alecian et al. 2011), HD 121743 (Briquet et al. in preparation), HD 127381 (Henrichs et al. 2012), HD 130807 (Buysschaert et al. 2017), HD 136504 (Shultz et al. 2015b; Pablo et al. in preparation), HD 142184 (Grunhut et al. 2012a), HD 149438 (Kochukhov et al. in preparation), HD 156324 (Shultz et al. 2018), HD 163472 (Neiner et al. 2012a), HD 164492C (Wade et al. 2017), HD 176582 (Neiner et al., in preparation), HD 182180 (Oksala et al. 2010; Rivinius et al. 2013), HD 184927 (Yakunin et al. 2015), and HD 205021 (Neiner et al., in preparation).

Measurements from ESPaDOnS, Narval, and HARPSpol are plotted with separate symbols and abbreviated in the legends as ESP, Nar, and HAR, respectively. The excellent agreement between ESPaDOnS and Narval measurements has been demonstrated by Wade et al. (2016). Those stars for which ESPaDOnS and HARPSpol measurements of similar quality are both available (e.g. HD 130807, HD 133880, HD 164492C; Buysschaert et al. 2017; Kochukhov et al. 2017; Wade et al. 2017) demonstrate that longitudinal magnetic field measurements collected with these instruments also agree within 1σ , a result which we confirm here for the cases of HD 64740, HD 67621, and HD 156424.

For comparisons to $\langle B_z \rangle$ measurements from the literature, the following abbreviations for the references are used in the legends: BL77 (Borra & Landstreet 1977); BL79 (Borra & Landstreet 1979); BLT83 (Borra et al. 1983); B87 (Bohlender et al. 1987); BLT93 (Bohlender, Landstreet & Thompson 1993); M91 (Mathys 1991); MH97 (Mathys & Hubrig 1997); B04 (Briquet et al. 2004); B06 (Bagnulo et al. 2006); B07 (Briquet et al. 2007); Y11 (Yakunin et al. 2011); BM11 (Bohlender & Monin 2011); Y13 (Yakunin 2013); B15 (Bagnulo et al. 2015); and F15 (Fossati et al. 2015b). Photopolarimetric $\langle B_z \rangle$ measurements conducted using the wings of the $H\beta$ line (BL77, BL79, BLT83, and B87) have been corrected to 80 per cent of the published values as suggested by Mathys et al. (2000) (see Section 3.2.2).

HD 25558: Sódor et al. (2014) determined that the rotational period of the magnetic secondary of this SB2 system should be about 1.2 ± 0.6 d. Using the restricted data set described in Section 3.2,

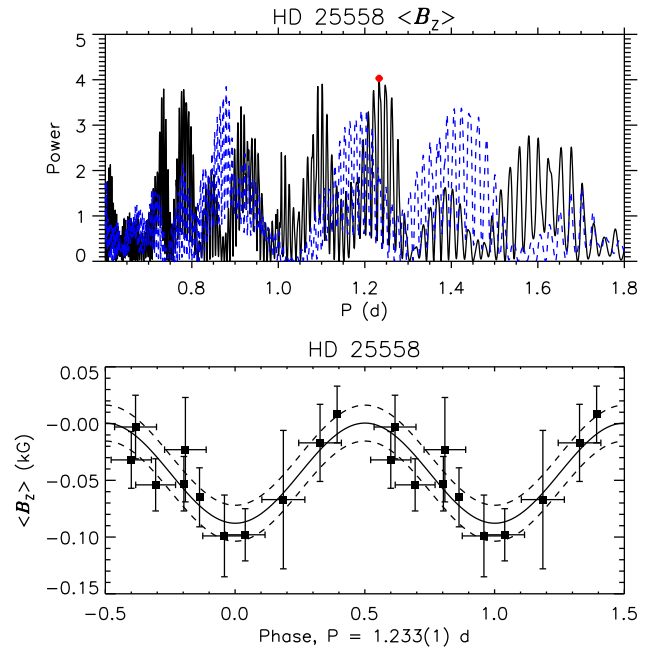


Figure A1. Top: periodogram for $\langle B_z \rangle$ measurements of HD 25558. The abscissa is limited to the period window. Solid (black) line indicates the periodogram for $\langle B_z \rangle$; dashed (blue) line the periodogram for $\langle N_z \rangle$. The red circle indicates the adopted rotation period. Bottom: $\langle B_z \rangle$ phased with P_{rot} . Solid lines indicate the best-fitting sinusoidal curves; dashed lines, the 1σ uncertainty in the fit. In this and in other figures, the best-fitting sinusoidal curve is unless otherwise stated of first order.

and restricting the period window to the range given by Sódor et al. (2014), we find maximum amplitude at $1.233(1)$ d, although there are numerous nearby peaks that provide a comparable phasing of the data (Fig. A1, top). The FAP of the maximum amplitude peak in the $\langle B_z \rangle$ period spectrum is 0.12, only slightly lower than the maximum FAP in the $\langle N_z \rangle$ spectrum, 0.17. $\langle B_z \rangle$ is shown phased with this period in the bottom panel of Fig. A1.

HD 35298: North (1984) found a photometric period of $P_{\text{rot}} = 1.85336(1)$ d. The periodogram constructed from archival *Hipparcos* photometry shows maximum power at ~ 0.927 d, which is almost exactly half of the period reported by North (1984) (Fig. A2, top). This period does not provide an acceptable phasing of the magnetic data, however. The second-highest peak in the *Hipparcos* photometry periodogram is at $1.8548(2)$ d. The FAP of this peak is $\sim 10^{-3}$, much lower than the FAP of the highest peak in the null spectrum, ~ 0.999 . While the FAP of the highest peak in the ESPaDOnS $\langle B_z \rangle$ periodogram is only 0.03, maximum power in the $\langle B_z \rangle$ periodogram is at $1.8546(1)$ d, with no power in the $\langle N_z \rangle$ periodogram at this period. This period also provides an acceptable phasing of the dimaPol $H\beta$ $\langle B_z \rangle$ measurements. Including the best dimaPol measurements (i.e. those that do not fall outside the range of the ESPaDOnS measurements), along with the $\langle B_z \rangle$ measurements published by Yakunin (2013), increases the precision of the period to $1.85459(3)$ d, with an FAP of 10^{-8} . This is compatible with $P_{\text{rot}} = 1.85457$ d found by Yakunin (2013). $\langle B_z \rangle$ and H_p are shown phased with this period in Fig. A2. H_p shows a double-wave variation, accounting for the appearance of the 0.927 d peak. The $\langle B_z \rangle$ curve is fit best with a third-order sinusoid, which may be indicative of departures from a pure centred dipole.

HD 36485: $\langle B_z \rangle$ measurements of HD 36485 (δ Ori C) have been published by Bohlender et al. (1987), Mathys & Hubrig (1997),

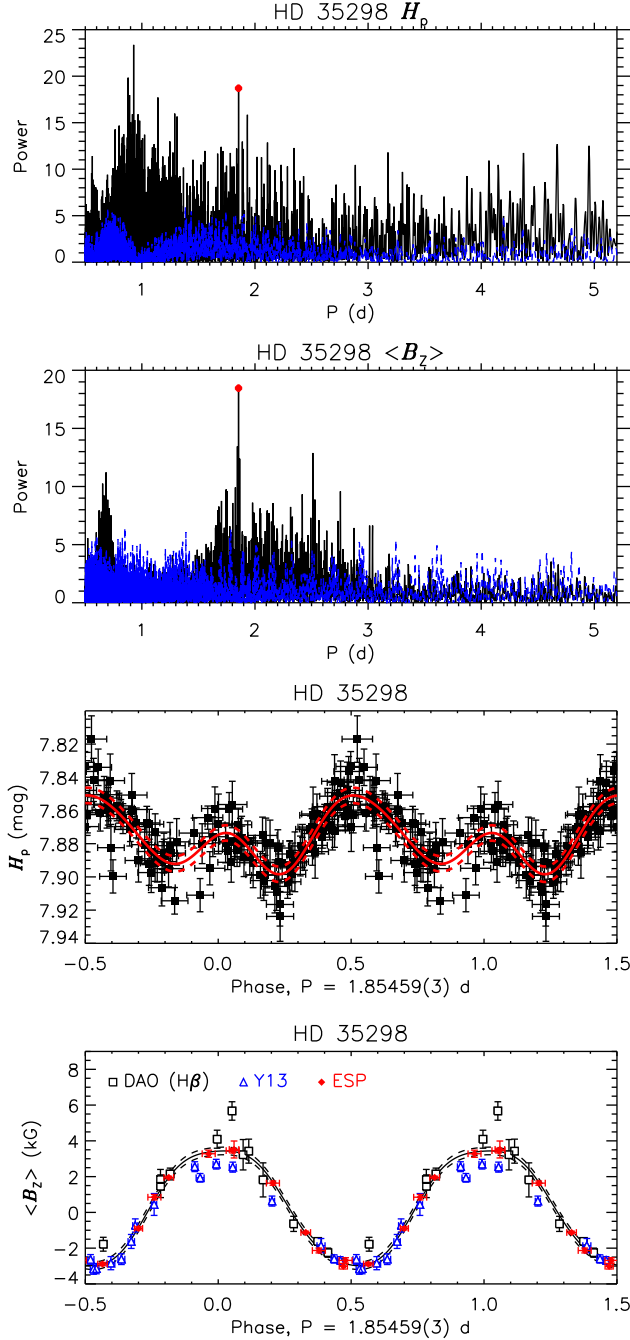


Figure A2. Rotational period of HD 35298. Top: periodogram constructed for H_p photometry (solid black line) and synthetic null measurements (dashed blue line). The adopted period is indicated by the red circle. Second panel: periodogram from ESPaDOnS $\langle B_z \rangle$ measurements. Third panel: H_p phased with P_{rot} . The red curve is second-order sinusoidal fit; dashed curves indicate uncertainties in the fit. Bottom: $\langle B_z \rangle$ phased with P_{rot} . The sinusoidal fit is of third order.

Leone et al. (2010), and Yakunin et al. (2011). ESPaDOnS data are unable to improve the period, which was determined spectroscopically using the magnetic primary’s variable $H\alpha$ emission (Leone et al. 2010). They are consistent in magnitude with previous data (see Fig. A3), but much more precise: in contrast to earlier measurements, a weak variation with a semi-amplitude of 0.06 ± 0.03 kG can be discerned in the high-resolution measurements. The ESPaDOnS data agree in magnitude with the Special Astrophysical

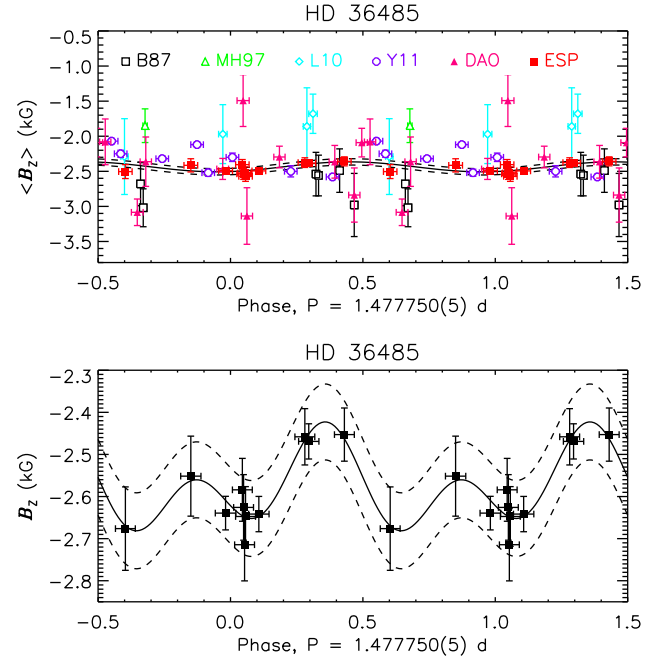


Figure A3. $\langle B_z \rangle$ curves for HD 36485 (δ Ori C). The top panel shows a comparison to historical data, demonstrating the much higher precision of the ESPaDOnS measurements. The bottom panel shows ESPaDOnS measurements only. Note that the $\langle B_z \rangle$ curve is apparently better fit by a second-order sinusoid.

Observatory observations reported by Yakunin et al. (2011), but there is less scatter in the new data. The bottom panel shows the ESPaDOnS measurements only, fitted with a second-order sinusoid, which provides a somewhat better fit to the data. The possibility that the star’s magnetic field contains significant quadrupolar contributions was earlier suggested on the basis of a pronounced asymmetry in the magnetospheric $H\alpha$ emission (Leone et al. 2010). While the ESPaDOnS measurements are compatible with this hypothesis, and indeed represent the best test of it to date, the extremely low level of variation means that it cannot yet be unambiguously confirmed. Measurements with a precision of approximately $3\times$ that of the existing data set should be sufficient to definitively test the hypothesis that HD 36485 possesses a complex surface magnetic field.

HD36526: North (1984) found a photometric period of 1.5405(1) d. This period provides a somewhat imperfect phasing of the ESPaDOnS $\langle B_z \rangle$ measurements; period analysis of the new ESPaDOnS magnetic data indicate a slightly longer $P_{\text{rot}} = 1.542(1)$ d, while the dimaPol $H\beta$ (B_z) measurements yield 1.5419(2) d. Combining the data sets yields 1.54185(4) d (Fig. A4). The FAP of this peak is 2×10^{-4} , while the highest peak in the $\langle N_z \rangle$ periodogram has an FAP of 0.01. (B_z) is shown phased with this period in the bottom panel of Fig. A4.

HD 36982: spectroscopic variability is likely associated with this Herbig Be star’s accretion disc, and therefore cannot be used to constrain P_{rot} . Within the period window, the maximum power in the $\langle B_z \rangle$ periodogram is at 0.425 d. However, there is a strong peak in the $\langle N_z \rangle$ periodogram at this period (Fig. A5, top), and this is furthermore below the Nyquist frequency of 0.8 d^{-1} . We therefore adopt the next-strongest peak, 1.8551(5) d. However, it should be noted that the FAPs of the highest peaks in $\langle B_z \rangle$ and $\langle N_z \rangle$ are both similar, about 0.03. (B_z) is shown phased with this period in Fig. A5 (bottom).

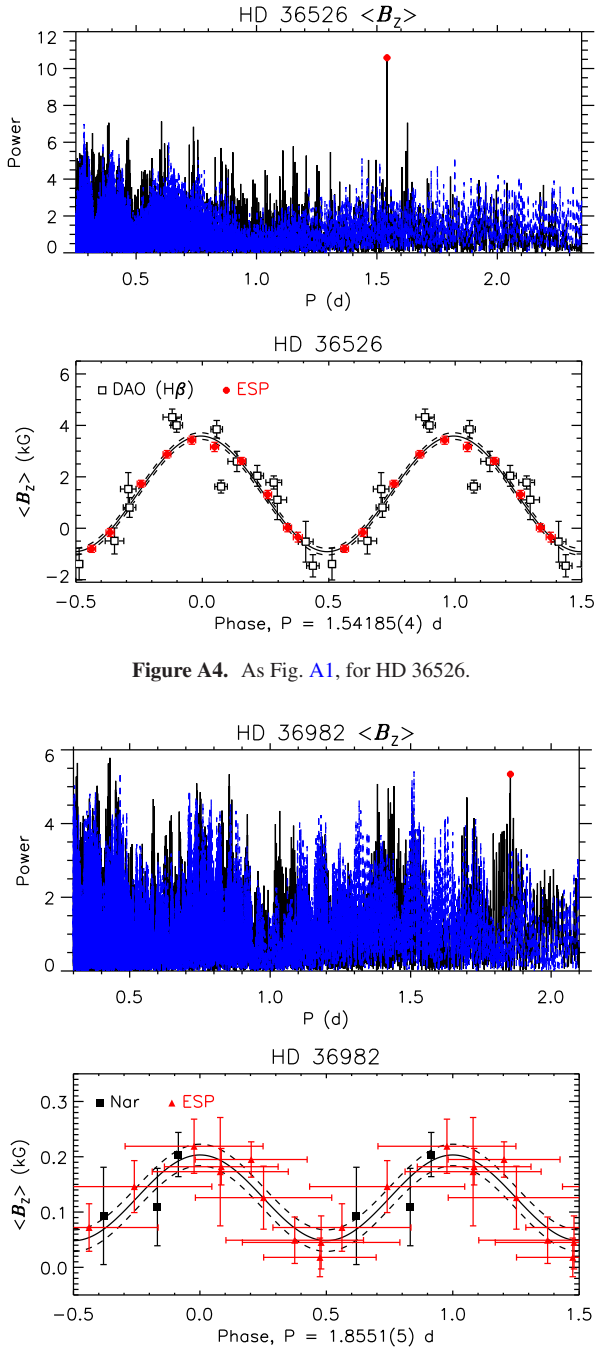
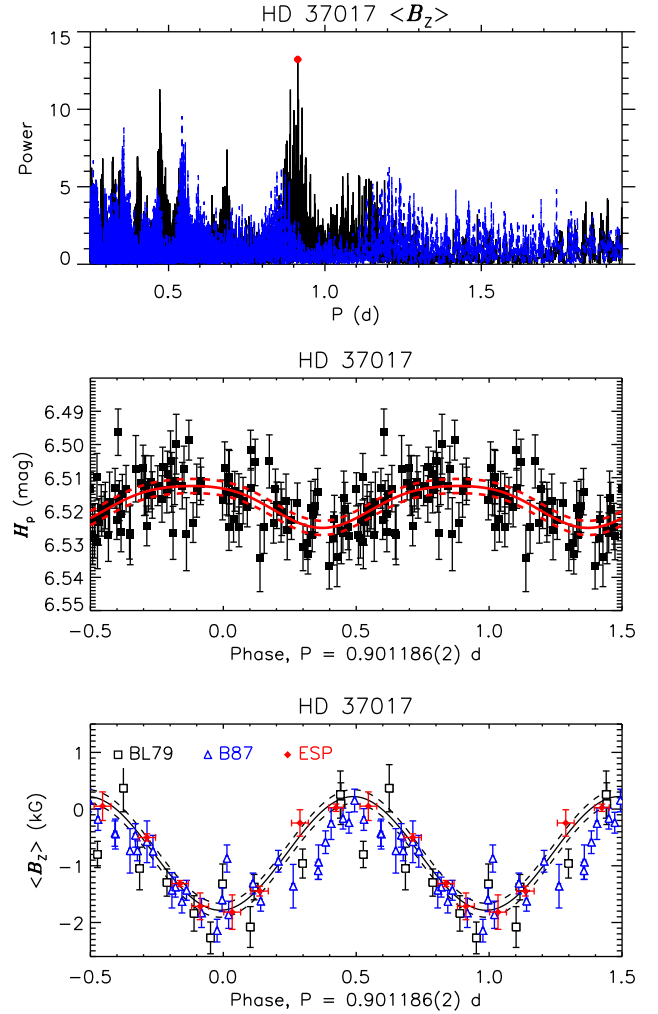


Figure A4. As Fig. A1, for HD 36526.

Figure A5. As Fig. A1, for HD 36982.

HD 37017: Bohlender et al. (1987) combined periodograms for photometric and magnetic data and found $P_{\text{rot}} = 0.901195(5)$ d. This period does not provide a reasonable phasing of the ESPaDONs data. Combining historical measurements with our own, we find $0.901186(2)$ d. The maximum amplitude period has an FAP of about 10^{-7} , while the FAP of the highest amplitude peak in the $\langle N_z \rangle$ spectrum is 0.12. The periodogram and $\langle B_z \rangle$ measurements are shown in the top and middle panels of Fig. A6. The *Hipparcos* photometry is also coherently phased with this period (bottom panel), although the low level of photometric variation means that it cannot be used to determine a more precise period.

HD 37058: Pedersen (1979) reported a photometric period of ~ 14 d, while Bychkov et al. (2005) determined a 1.022 d period. In

Figure A6. Top: $\langle B_z \rangle$ periodogram for HD 37017. Middle: *Hipparcos* photometry phased with P_{rot} . The red curve is second-order sinusoidal fit; dashed curves indicate uncertainties in the fit. Bottom: $\langle B_z \rangle$ phased with P_{rot} .

both cases, the available data sets were too small to firmly establish P_{rot} . The ESPaDONs measurements confirm Pedersen's (1979) period. By combining with $\langle B_z \rangle$ measurements from the literature, we find $P_{\text{rot}} = 14.581(2)$ d, with an FAP of 0.03, much lower than the FAP of the maximum amplitude peak in the null spectrum, ~ 0.3 . $\langle B_z \rangle$ is shown phased with this period in Fig. A7.

HD 37061: while HD 37061 (NU Ori) displays spectroscopic variability, it is not clear that this is purely due to rotational modulation as periodograms created for these data yield contradictory results. Therefore, we base our period search on the magnetic data. The star's high $v \sin i$ ($225 \pm 8 \text{ km s}^{-1}$) means the period must be no longer than ~ 1.4 d. There are two peaks of comparable amplitude within the period window, at ~ 0.6 and ~ 1.1 d (Fig. A8, top). The shortest period would indicate $v_{\text{eq}} \sim v_{\text{br}}$, necessitating very tight limits on the stellar parameters. The FAPs of either peak are both about 0.08, only slightly below the $\langle N_z \rangle$ maximum amplitude FAP of about 0.1. As the next-highest peak yields less extreme rotational parameters, we adopt the peak at 1.0950(4) d, as the more conservative option. $\langle B_z \rangle$ is shown phased with this period in Fig. A8.

HD 55522: Briquet et al. (2004) used spectroscopy and ground-based as well as archival *Hipparcos* photometry to find $P_{\text{rot}} = 2.729(1)$ d. This period phases the ESPaDONs data well

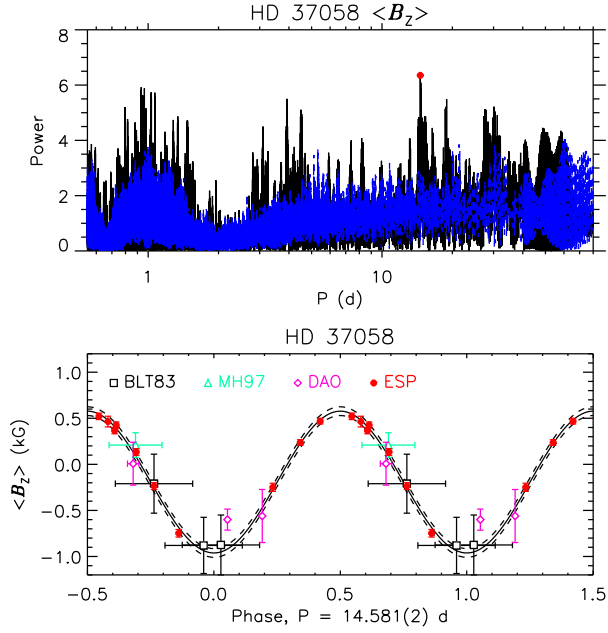


Figure A7. As Fig. A1, for HD 37058.

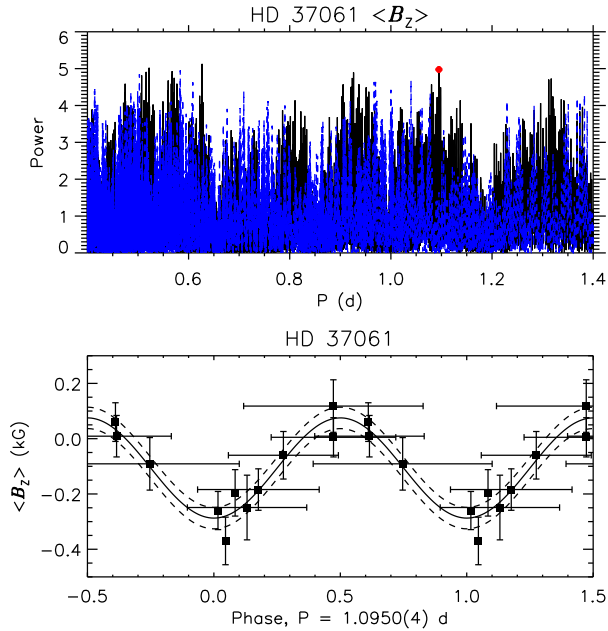


Figure A8. As Fig. A1, for HD 37061.

when examined in isolation, but comparison of the $\langle B_z \rangle$ measurements presented by Briquet et al. (2007) to our own reveals a small phase offset. While there are numerous closely spaced peaks in the periodogram constructed for the combined magnetic data, the only peak consistent with the photometric data is at 2.7292(3) d. $\langle B_z \rangle$ and H_p are shown phased with this period in Fig. A9. While the phase uncertainties are somewhat large, maximum light appears to occur at about $\langle B_z \rangle \sim 0$.

HD 58260: Pedersen (1979) suggested $P_{\text{rot}} = 1.657$ d based on the star's photometric variation. While this period is compatible with the period window of 1.42–164 d, we cannot confirm this period using *Hipparcos* photometry. Bohlender et al. (1987) were unable to determine a period for this star due to the negligible variation in

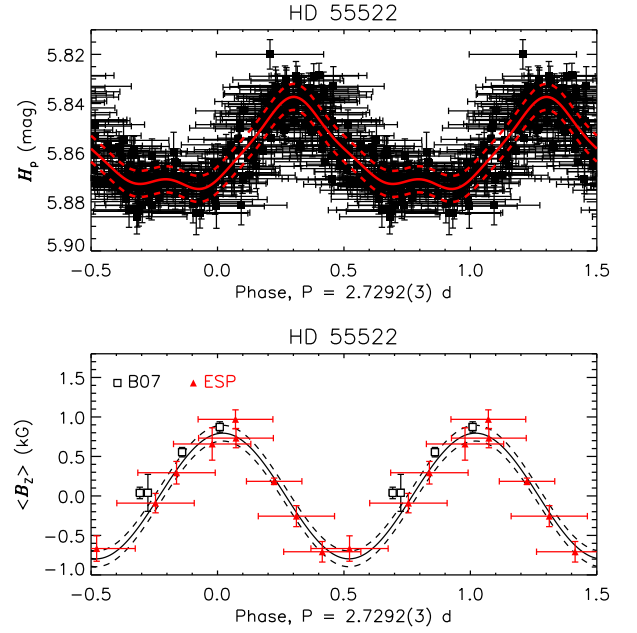


Figure A9. Top: *Hipparcos* photometry phased with the revised period for HD 55522. The red curve is a second-order sinusoidal fit; and dashed curves indicate the uncertainty in the fit. Bottom: $\langle B_z \rangle$, phased with the revised period. The solid curve shows the best-fitting sinusoid; and the dashed curve shows the 1σ uncertainty.

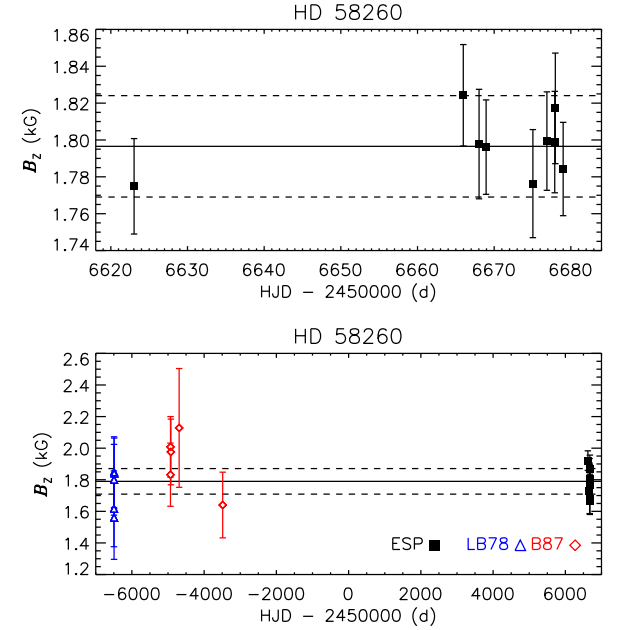


Figure A10. $\langle B_z \rangle$ measurements of HD 58260 as a function of time. The top panel shows the ESPaDOnS data, and the bottom panel a comparison to historical data. Solid lines indicate the mean ESPaDOnS $\langle B_z \rangle$ and dashed lines the mean ESPaDOnS $\langle B_z \rangle$ error bar.

$\langle B_z \rangle$. Even at the much higher precision of the ESPaDOnS data, the standard deviation of $\langle B_z \rangle$ is only 17 G, with a mean error bar of 28 G (Fig. A10, top). While there are some higher amplitude peaks at short periods, and very little power at periods longer than ~ 10 d, this is likely an artefact caused by the relatively short temporal baseline (56 d, with all but 1 observation acquired within 13 d). Analysis of *Hipparcos* photometry yields no peaks with an S/N in

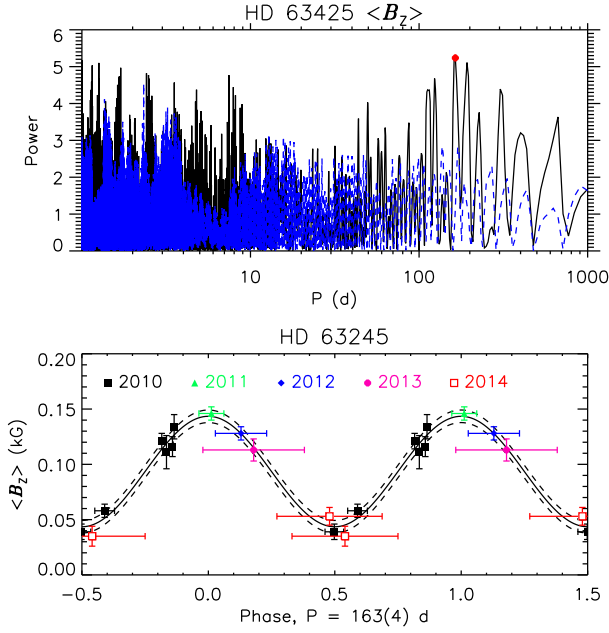


Figure A11. As Fig. A1, for HD 63425.

excess of 4 at periods of greater than 1 d. Comparison to historical data (Fig. A10, bottom panel) demonstrates that $\langle B_z \rangle$ is consistent with no variation over a time-scale of ~ 35 yr: all measurements are within 1σ of the mean ESPaDOnS $\langle B_z \rangle$.

HD 63425: two periods are compatible with the existing $\langle B_z \rangle$ measurements. Maximum power is at 163(4) d, which is compatible with the low $v \sin i$. The other is very short, ~ 0.55 d, which would make HD 63425 amongst the most rapidly rotating magnetic massive stars. The FAPs of the two periods are similar, about 0.08, and only slightly lower than the minimum $\langle N_z \rangle$ FAP of 0.1. Given the very small $v \sin i$, this would require the rotational pole to be almost perfectly aligned with the line of sight, which is a priori less likely. Furthermore, the $\langle B_z \rangle$ measurements from 2010 (the year with the largest number of observations) show statistically significant differences only between those measurements separated by ~ 250 d, while the variation of those collected within ~ 10 d is of the same order as the error bars. We therefore adopt the longer period. $\langle B_z \rangle$ is shown phased with this period in Fig. A11.

HD 64740: combining photometric and magnetic data, Bohlender et al. (1987) found 1.33026(6) d. Combining the magnetic measurements presented by Borra & Landstreet (1979) and Bohlender et al. (1987) with the new ESPaDOnS and HARPSpol data, we find $P_{\text{rot}} = 1.330205(3)$ d, with an FAP of 4×10^{-7} , much lower than the minimum $\langle N_z \rangle$ FAP of 0.28. The periodogram, $\langle B_z \rangle$ measurements, and H_p light curve are shown in Fig. A12. Despite the relatively strong magnetic field, the light curve is only weakly variable, and could not be used to constrain P_{rot} . A second-order sinusoidal fit has been used for $\langle B_z \rangle$, which may indicate that the surface magnetic field is slightly more complex than a centred dipole.

HD 66522: a period cannot be determined uniquely from the sparse magnetic data, but the *Hipparcos* photometry shows a clear modulation at 909(75) d. The FAP is 2×10^{-14} , while the minimum FAP in the null period spectrum is 0.99. This extremely long period is compatible with the star's negligible $v \sin i$. The H_p periodogram is shown in the top panel of Fig. A13; H_p and $\langle B_z \rangle$ are shown phased with the 909 d period in the middle and bottom panels. We chose the date of maximum light as JD0, which seems to correlate well

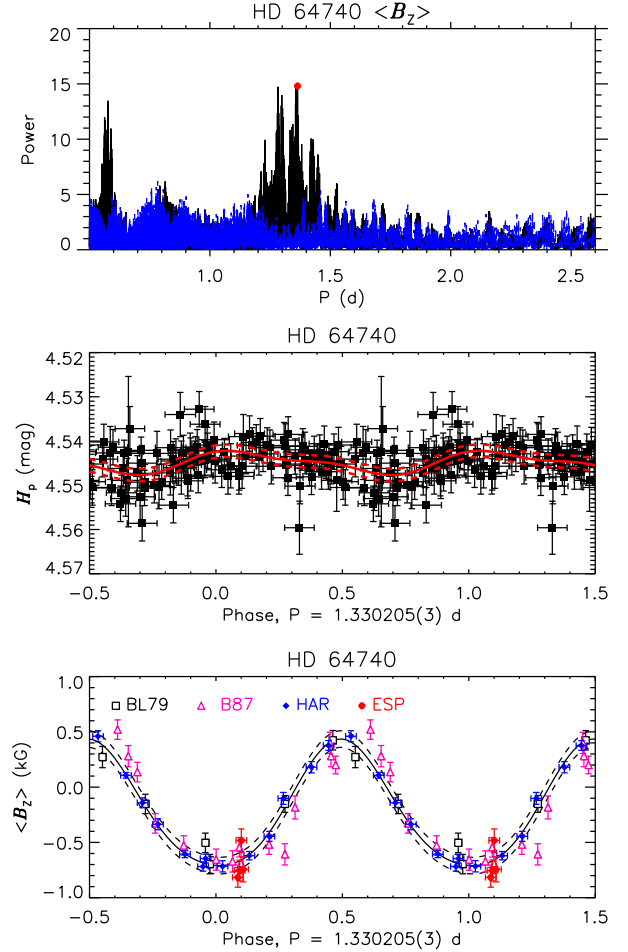


Figure A12. As Fig. A6 for HD 64740. The sinusoidal fit is of second order.

to $|\langle B_z \rangle|_{\text{max}}$ (although it is worth noting that the phase uncertainties are very large for the magnetic data, about 0.8 cycles). The star also shows spectroscopic variation; minimum line strength corresponds to maximum light, and vice versa, suggesting that both spectroscopic and photometric variations are due to the rotational modulation of chemical spots.

HD 66665: Petit et al. (2013) reported a period of 21 d based upon a preliminary analysis of ESPaDOnS and Narval magnetic measurements. We find 24.5(1) d using the same data, with an FAP of 2×10^{-4} . The 21 d period is definitely ruled out. There is very little power in the $\langle N_z \rangle$ periodogram near 25 d, and the minimum FAP of the $\langle N_z \rangle$ period spectrum is 0.08. The periodogram is shown in the top panel of Fig. A14; $\langle B_z \rangle$ is shown phased with this period in the bottom panel.

HD 66765: Alecian et al. (2014) found $P_{\text{rot}} = 1.62(15)$ d based on 4 d of observations with HARPSpol. We acquired a single new ESPaDOnS observation approximately 2 yr later. The quality of the measurement is not as high as that of the HARPSpol data, and does not enable improvement of the period. However, we also acquired 11 FEROS spectra. As HD 66765 is a He-variable star, we measured the EWs of the He I 447.1, 587.6, and 667.8 nm lines. Within the uncertainty in P_{rot} determined by Alecian et al. (2011), the highest peak in the periodogram for the combined EWs is at 1.6079(5) d, with an FAP of 2×10^{-4} and a minimum FAP in the null spectrum of 0.58. While the star is only weakly photometrically variable, the highest peak in the *Hipparcos* periodogram is at 1.6079(2) d, some-

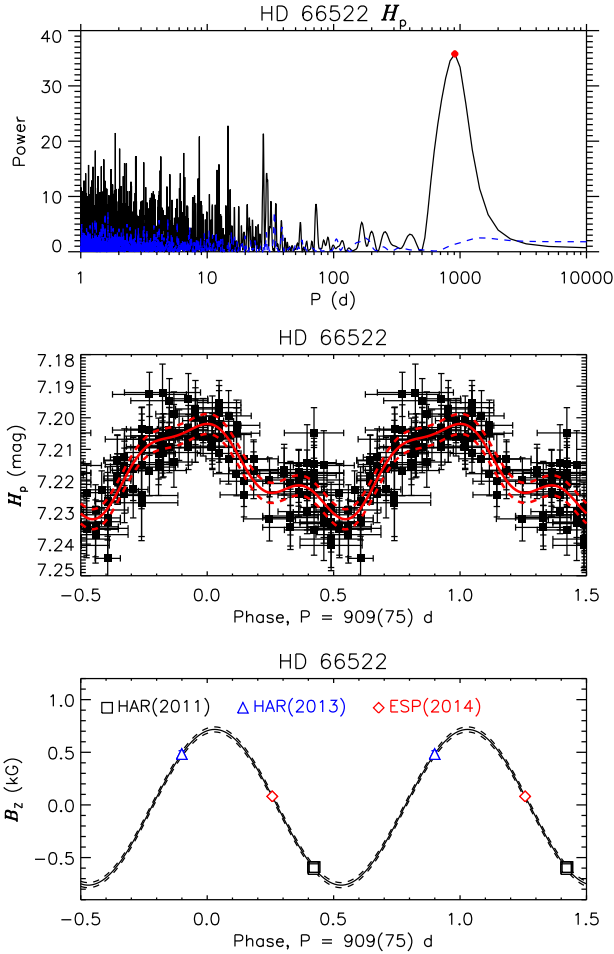


Figure A13. Top: *Hipparcos* photometry periodogram for HD 66522. Middle: H_p light curve. Bottom: $\langle B_z \rangle$ curve. Numbers in brackets in the legend indicate the year of observation. In the bottom panel, phase errors are larger than the x -axis and are not shown.

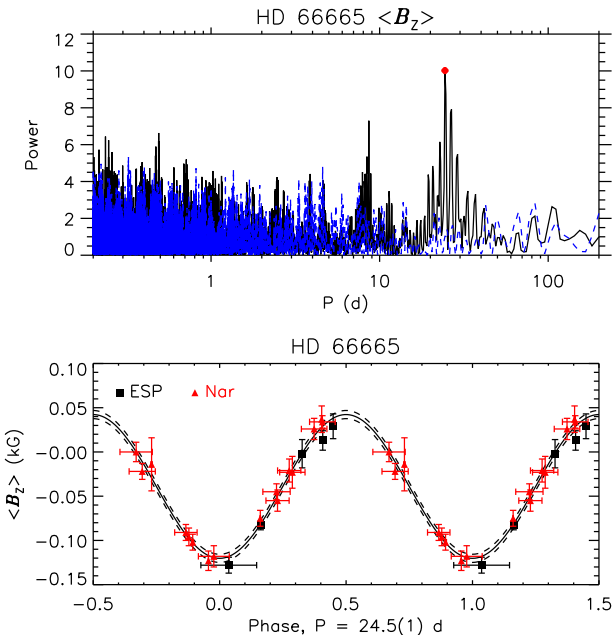


Figure A14. As Fig. A1, for HD 66665.

what more precise than the spectroscopic determination, although the FAP is only 0.54 (Fig. A15, top). The combined He I EWs, H_p , and the $\langle B_z \rangle$ measurements are shown phased with this period in the bottom three panels of Fig. A15. Note that phase uncertainties are not shown for the light curve as, with the precision of this period, they are larger than the x -axis.

HD 67621: Alecian et al. (2014) determined $P_{\text{rot}} = 3.6(2)$ d. As with HD 66765, we observed the star once more with ESPaDOnS, and refine the period to 3.593(1) d, although we note that the periodogram has numerous nearby peaks within the original range (Fig. A16, top). While the minimum FAPs of the $\langle B_z \rangle$ and $\langle N_z \rangle$ period spectra are similar (0.27 and 0.31, respectively), there is no power near the 3.6 d period in the $\langle N_z \rangle$ periodogram. $\langle B_z \rangle$ is shown phased with this period in Fig. A16 (bottom). We attempted to constrain the period using *Hipparcos* photometry, but the almost complete absence of variability means that the light curve is dominated by noise. This is consistent with the star's He lines, which also show very low levels of variability.

HD 122451: in order to maximize the S/N, we used $\langle B_z \rangle$ measurements from LSD profiles in 1 d bins. The lower bound of the period window as defined by the break-up velocity, ~ 0.5 d, is at a higher frequency than the Nyquist frequency of $\sim 0.5 \text{ d}^{-1}$, therefore the Nyquist frequency was taken as the lower bound. The highest peak in the $\langle B_z \rangle$ periodogram is at 2.885(1) d (Fig. A17, top). $\langle B_z \rangle$ is shown phased with this period in Fig. A17 (bottom). While this period achieves an S/N of 5.4 and is thus formally significant, it must be noted that the FAP is 0.06, which is close to the minimum FAP in the $\langle N_z \rangle$ spectrum, 0.09, so this period should be regarded as tentative. However, it is close to a period identified by Pigulski et al. (2016) in the BRITE light curve as a possible rotational period, 2.827(2) d.

HD 125823: the most recent ephemeris for HD 125823 was given by Catalano & Leone (1996), who found a period of 8.8177(1) d by combining photometric data with an approximately 20 yr baseline. This period phases the ESPaDOnS data well if examined in isolation, but there is a systematic phase offset when compared to older data. The highest amplitude in the periodogram for the ESPaDOnS data is at 8.817(4) d. This period phases the ESPaDOnS $\langle B_z \rangle$ measurements well with the historical data; a periodogram constructed from the combined data sets yields a more precise period, $P_{\text{rot}} = 8.8169(1)$ d. The FAP of this peak is 2.2×10^{-5} , much lower than the minimum $\langle N_z \rangle$ FAP of 0.33. This period also provides an acceptable phasing of the FORS2 measurements reported by Fossati et al. (2015b), who noted they were unable to phase their measurements with the historical data using the Catalano & Leone (1996) ephemeris. $\langle B_z \rangle$ and H_p are shown phased with this period in Fig. A18. The ephemeris results in the photometric and magnetic extrema coinciding. This is consistent with the strong He-line variability, which changes from He-strong at the positive magnetic pole to He-weak at the negative magnetic pole.

HD 142990: Bychkov et al. (2005) found a rotation period of 0.97907 d. The magnetic data, even including the new measurements presented here, are rather sparse and cannot be used to identify a unique period within the period window. However, a unique period of 0.97903(5) d is clearly identified in the periodogram obtained using archival *Hipparcos* photometry, suggesting the Bychkov et al. period is indeed correct (Fig. A19, top). The FAP of this peak is 2.3×10^{-17} , much lower than the minimum FAP in the null period spectrum of 0.5. This period provides an adequate phasing of the ESPaDOnS $\langle B_z \rangle$ measurements, but is insufficiently precise to phase the new data with the measurements reported by Borra et al. (1983) and Bohlender et al. (1993). A periodogram con-

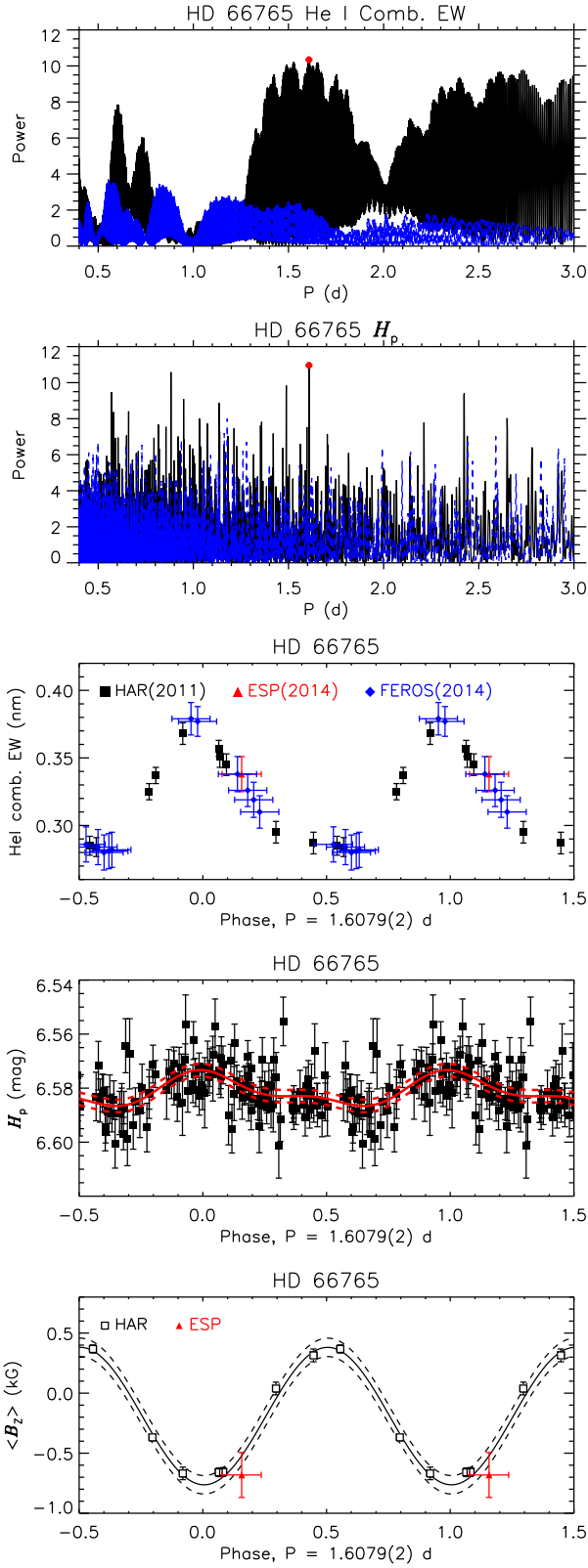


Figure A15. Top–bottom: periodogram for combined He I EWs; periodogram for *Hipparcos* photometry; combined He I EWs; *Hipparcos* photometry; $\langle B_z \rangle$ measurements. P_{rot} was determined by combining H_p and He I EW periodograms.

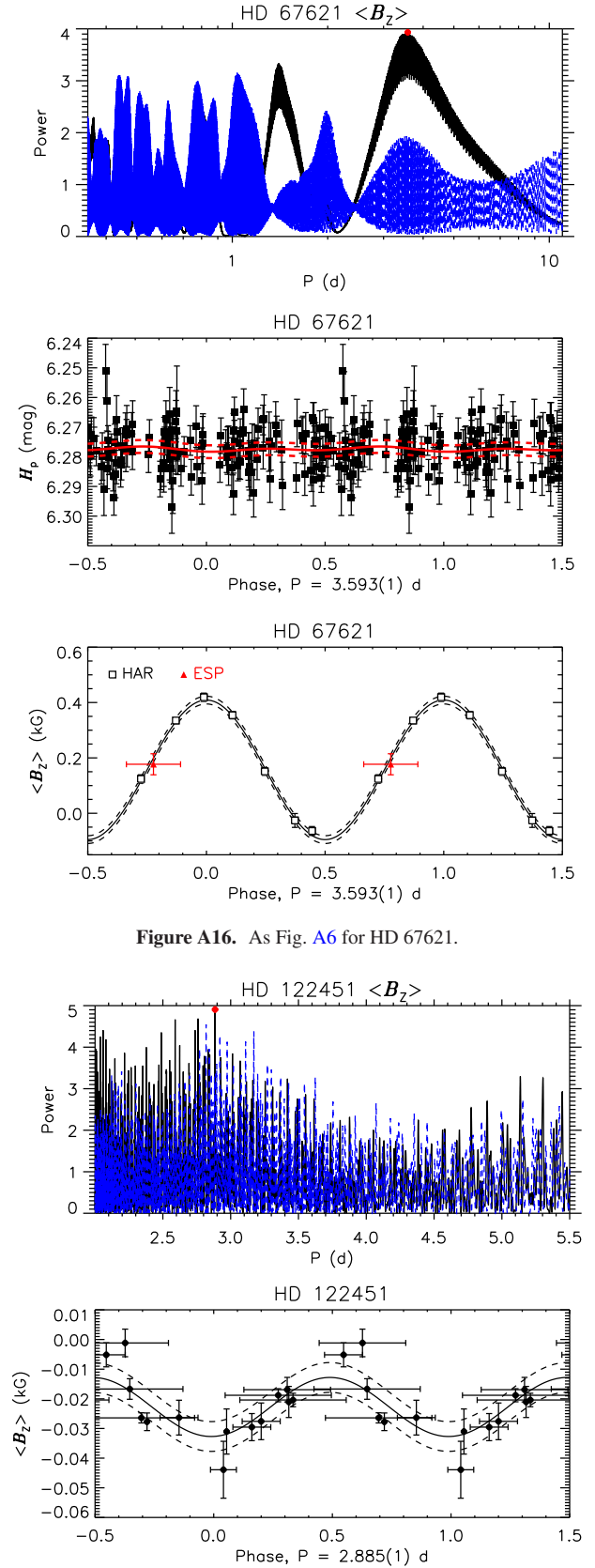


Figure A16. As Fig. A6 for HD 67621.

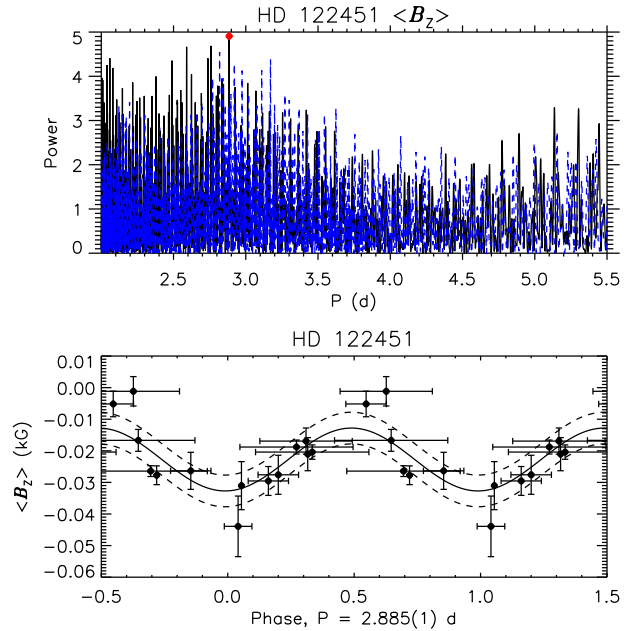


Figure A17. As Fig. A1, for HD 122451.

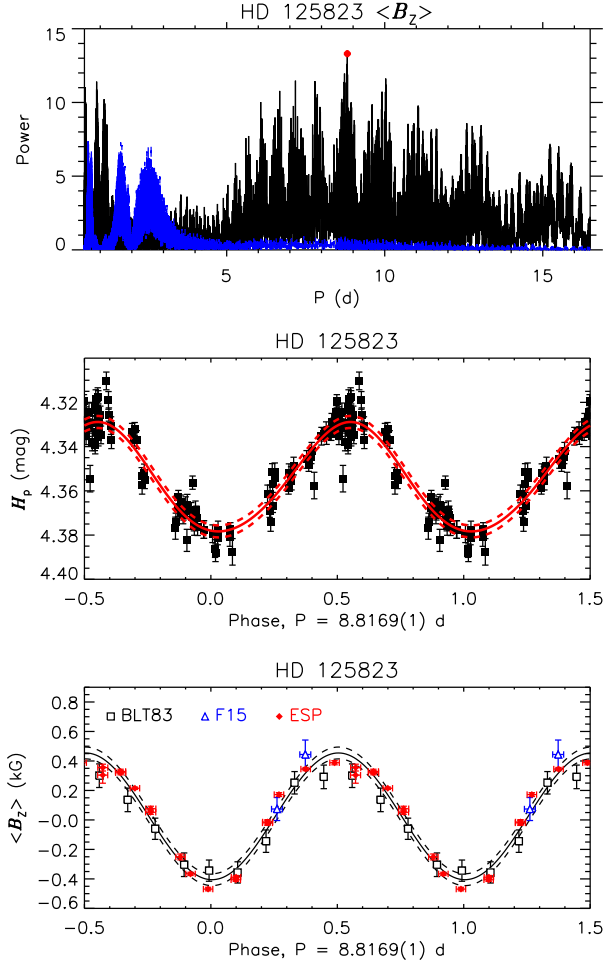


Figure A18. As Fig. A6, for HD 125823.

structured for all old and new $\langle B_z \rangle$ measurements, limited to a narrow window around 0.979 d, finds maximum amplitude at 0.978832(2) d, with an FAP of 0.0003 in $\langle B_z \rangle$ and 0.52 in $\langle N_z \rangle$. H_p and $\langle B_z \rangle$ are shown phased with this period in the middle and bottom panels of Fig. A19. Maximum light coincides with the negative magnetic pole. A second-order sinusoid has been used to fit $\langle B_z \rangle$, as this provides a better fit than a single-order sinusoid.

HD 149277: the strongest peak in the periodogram is at 25.380(7) d (Fig. A20, top), with an FAP of 9×10^{-5} , as compared to a minimum FAP in the $\langle N_z \rangle$ periodogram of 0.02. This period was determined at a lower precision using the ESPaDOnS data obtained before 2015, and with only slight modification successfully phases the newer ESPaDOnS data obtained in 2015, as well as the single FORS1 measurement reported by Bagnulo et al. (2006). We therefore consider this to be the correct rotation period with high confidence. $\langle B_z \rangle$ is shown phased with this period in Fig. A20 (bottom).

HD 156424: as this star shows weak emission in the high-velocity wings of its $H\alpha$ line, presumably originating in its magnetosphere, we conducted period analyses using both the magnetic data and $H\alpha$ EWs, with the EWs supplemented by FEROS spectroscopy. The variation in either case is small compared to the error bar. Almost all the $H\alpha$ variation is between data sets, suggesting that it is a consequence of systematic differences between instruments. There is very little power in either periodogram at periods longer than a few days, suggesting that the rotation period must be fairly

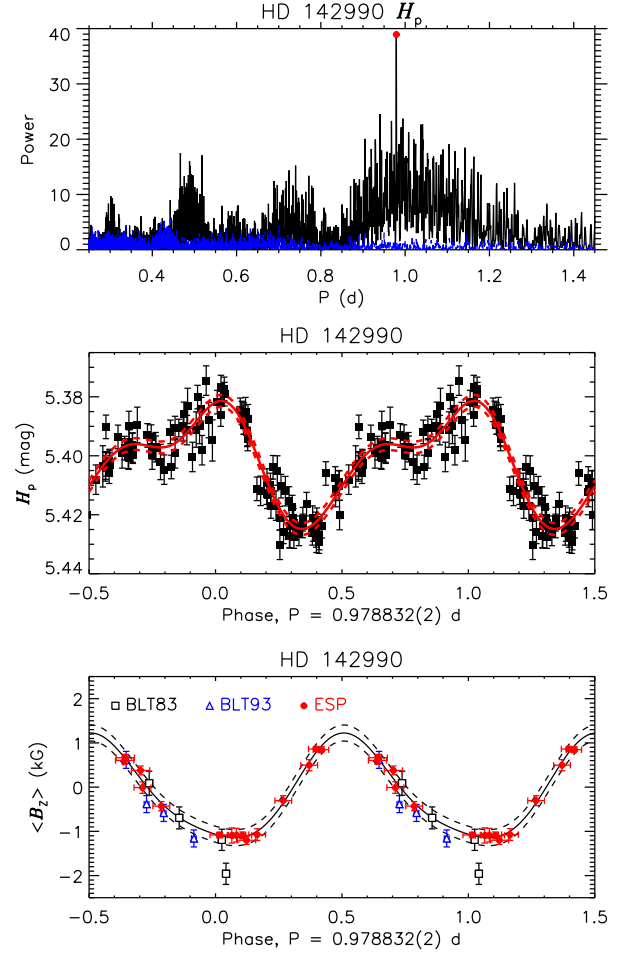


Figure A19. As Fig. A2, for HD 142990. The sinusoidal fit is of second order.

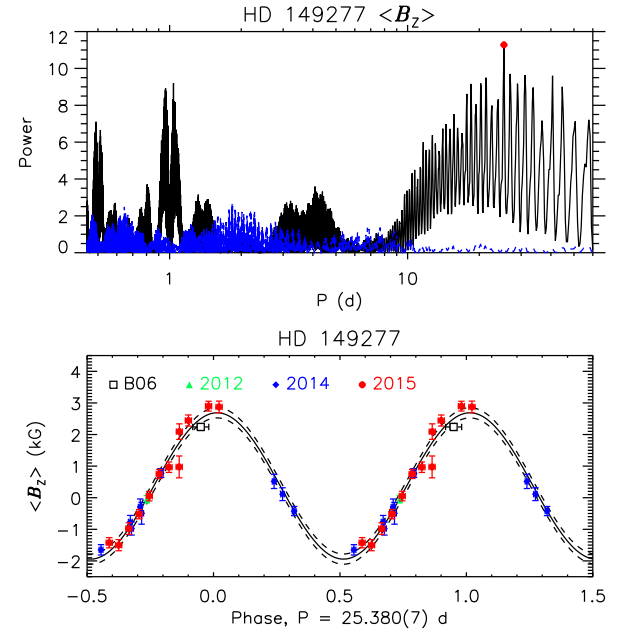


Figure A20. As Fig. A1, for HD 149277.

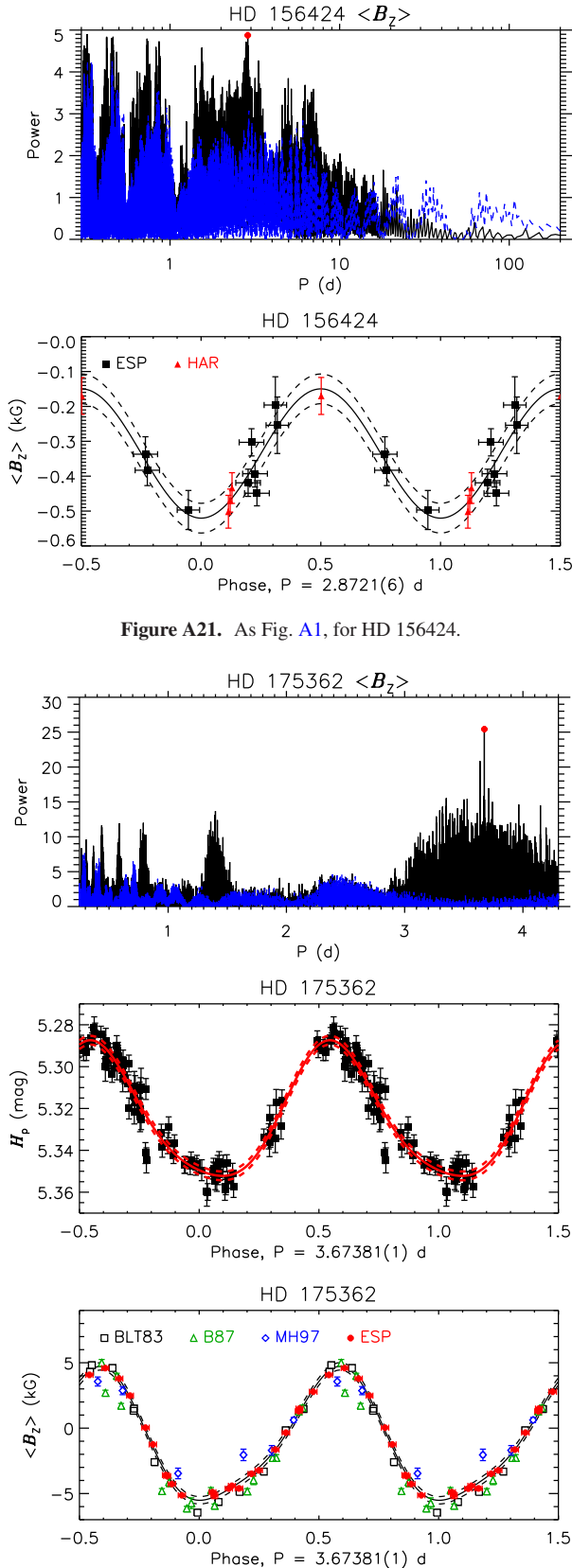


Figure A21. As Fig. A1, for HD 156424.

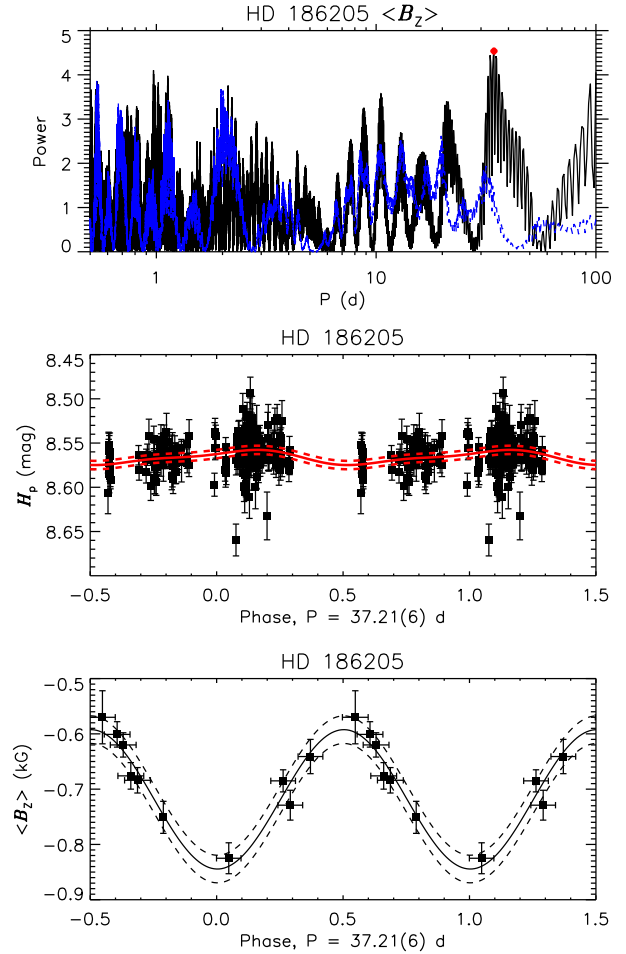


Figure A23. As Fig. A6, for HD 186205.

short despite the star's low $v \sin i$ (see Fig. A21, top). The largest amplitude peak in the $\langle B_z \rangle$ periodogram without a corresponding peak in the $\langle N_z \rangle$ periodogram is at 2.8721(6) d. $\langle B_z \rangle$ is shown phased with this period in the bottom panel of Fig. A21. It should be noted that the FAP of the highest peaks in the $\langle B_z \rangle$ period spectrum are around 0.06, not much lower than the minimum in $\langle N_z \rangle$, 0.15.

HD 175362: Bohlender et al. (1987) determined $P_{\text{rot}} = 3.6738(4)$ d for this star. Period analysis of the *Hipparcos* photometry yields a slightly less precise, but compatible period of 3.6732(6) d. Combining historical $\langle B_z \rangle$ measurements with ESPaDOnS $\langle B_z \rangle$ measurements enables us to refine the period to 3.67381(1) d (Fig. A22, top). $\langle B_z \rangle$ and H_p are shown phased with this period in the middle and bottom panels of Fig. A22. Maximum light occurs at the positive magnetic pole. This corresponds well to the He EWs, which are at a minimum at $\langle B_z \rangle_{\text{max}}$, and a maximum and $\langle B_z \rangle_{\text{min}}$.

HD 186205: the strongest peak in the $\langle B_z \rangle$ periodogram is at 37.21(6) d (Fig. A23, top), with an FAP of 0.12 (only slightly lower than the minimum FAP in $\langle N_z \rangle$, 0.20). There is very little power at either much higher or much lower periods, and the long period is consistent with the low $v \sin i$. Furthermore, there is no power in the $\langle N_z \rangle$ periodogram near this period. However, there are numerous nearby peaks which phase the data approximately as well. $\langle B_z \rangle$ is shown phased with this period in Fig. A23 (bottom). The semi-amplitude of the star's $\langle B_z \rangle$ variation, ~ 120 G, is much less than the median error bar in the diMaPol $\langle B_z \rangle$ measurements of 343 G,

Figure A22. As Fig. A6, for HD 175362. The sinusoidal fit to $\langle B_z \rangle$ is of third order.

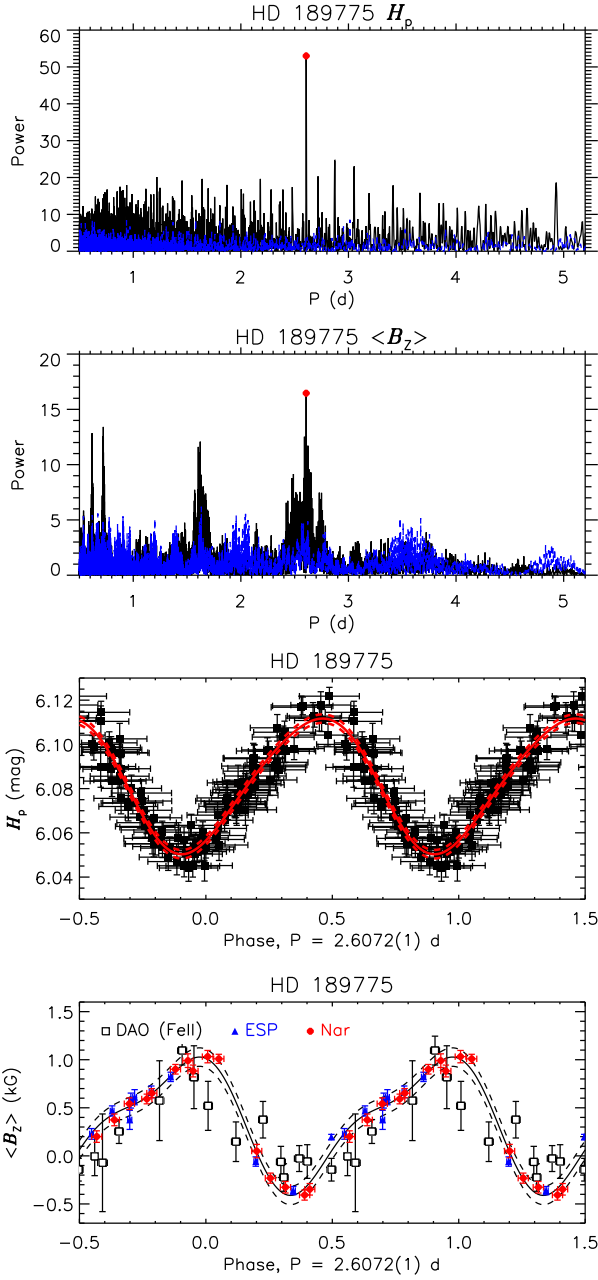


Figure A24. As Fig. A2, for HD 189775. The sinusoidal fit to $\langle B_z \rangle$ is of second order.

therefore the dimaPol data could not be used to help constrain P_{rot} . The star's He lines show very little variability; consistent with this, the *Hipparcos* light curve is dominated by noise (middle panel of Fig. A23), and also could not be used to constrain P_{rot} .

HD 189775: Petit et al. (2013) reported a period of 2.6048 d for this star, based on unpublished magnetic measurements collected with dimaPol. In addition to H_β $\langle B_z \rangle$ measurements, there are also $\langle B_z \rangle$ measurements conducted using the Fe II 492.3 nm line. As the H_β $\langle B_z \rangle$ measurements show significantly more scatter than the Fe II measurements, we adopt the latter. *Hipparcos* photometry shows a single peak at 2.6074(3) d. ESPaDOnS and Narval measurements yield a best-fitting period of 2.6071(3) d. For the combined magnetic data, we find 2.6072(1) d, as shown in the second panel of Fig. A24. The FAP of this peak in the periodogram is 3×10^{-7} ,

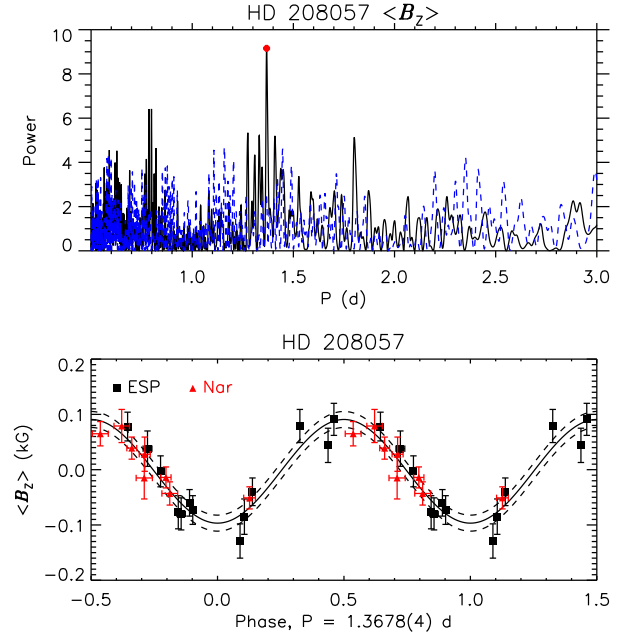


Figure A25. As Fig. A1, for HD 208057.

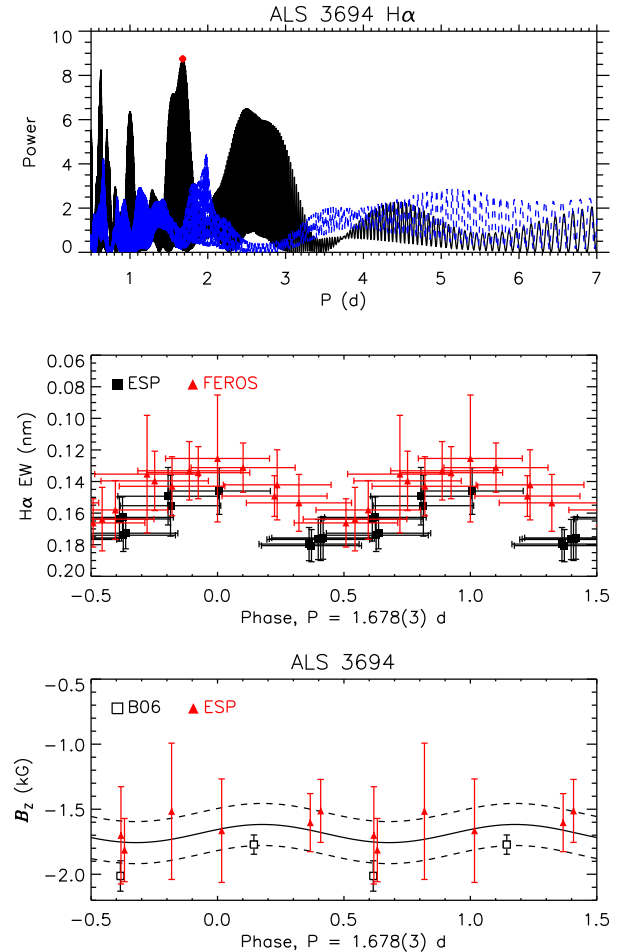


Figure A26. Top: periodogram for $H\alpha$ EWs of ALS 3694 (solid black line) and synthetic null measurements (dashed blue line). The adopted period is indicated by the red circle. Middle: $H\alpha$ EWs phased with P_{rot} . Bottom: $\langle B_z \rangle$ phased with P_{rot} .

lower than the FAP of the highest peak in $\langle N_z \rangle$, 0.1. The *Hipparcos* light curve and $\langle B_z \rangle$ measurements are shown phased with this period in Fig. A24. A second-order sinusoid was used to fit the magnetic data, as this star's $\langle B_z \rangle$ curve shows evidence of anharmonic behaviour that may be consistent with contributions from non-dipolar magnetic field components. Photometric extrema are somewhat offset from the magnetic extrema, suggesting that the chemical spots do not coincide precisely with the magnetic poles.

HD 208057: Henrichs et al. (2009) reported $P_{\text{rot}} = 1.441$ d. Using the same data, we find maximum power at 1.3678(4) d (Fig. A25, top). The light curve is dominated by pulsations, so the *Hipparcos* photometry could not be used to determine P_{rot} . However, it is worth noting that there is no power in the $\langle B_z \rangle$ periodogram at the pulsation periods of 1.23 and 1.25 d (Szewczuk & Daszyńska-Daszkiewicz 2015). The FAP of this peak is 7×10^{-4} , substantially smaller than the minimum FAP in the $\langle N_z \rangle$ period spectrum of 0.12. $\langle B_z \rangle$ is shown phased with this period in Fig. A25.

ALS 3694: no period can be determined from the magnetic data, as Σ_B is too low for any statistically significant variation to be detected.

However, this star displays $H\alpha$ emission consistent with an origin in a centrifugal magnetosphere (Shultz et al. 2016). Combining $H\alpha$ EWs measured from ESPaDOnS and FEROS spectra yields a periodogram with a single strong peak at 1.678(3) d (Fig. A26, top). The FAP of this peak is 0.001, much lower than the minimum FAP from the null periodogram, 0.43. $H\alpha$ EWs and $\langle B_z \rangle$ measurements are shown phased with this period in the middle and bottom panels of Fig. A26. Minimum $H\alpha$ EW (i.e. maximum emission strength) was used to define JD0, as this should correspond to maximum $\langle B_z \rangle$ (e.g. Townsend & Owocki 2005). While a sinusoidal variation is not distinguishable in $\langle B_z \rangle$, the magnitude of the ESPaDOnS $\langle B_z \rangle$ measurements is compatible with those published by Bagnulo et al. (2006).

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