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Abstract. We present a study of the composite supernova remnant G0.9+0.1 based on observations by *XMM-Newton*. The EPIC spectrum shows diffuse X-ray emission from the region corresponding to the radio shell. The X–ray spectrum of the whole Pulsar Wind Nebula is well fitted by an absorbed power–law model with a photon index $\Gamma \sim 1.9$ and a 2–10 keV luminosity of about $6.5 \times 10^{34}\ d_{10}^2$ erg s^{-1} (d_{10} is the distance in units of 10 kpc). However, there is a clear softening of the X-ray spectrum with distance from the core, which is most probably related to the finite lifetime of the synchrotron emitting electrons. This is fully consistent with the plerionic interpretation of the Pulsar Wind Nebula, in which an embedded pulsar injects energetic electrons into its surrounding region. At smaller scales, the eastern part of the arc-like feature, which was first revealed by *Chandra* observations, shows indications of a hard X-ray spectrum with a corresponding small photon index ($\Gamma = 1.0 \pm 0.7$), while the western part presents a significantly softer spectrum ($\Gamma = 3.2 \pm 0.7$). A possible explanation for this feature is fast rotation and subsequent Doppler boosting of electrons: the eastern part of the torus has a velocity component pointing towards the observer, while the western part has a velocity component in the opposite direction pointing away from the observer.

Key words. ISM: supernova remnants – ISM: individual objects: G0.9+0.1 – X–rays: ISM

1. Introduction

The radio source G0.9+0.1 is the only composite SNR, listed in the Green catalog (Green 2001), which lies in the general direction of the Galactic Center. According to Helfand & Becker (1987), the radio morphology of G0.9+0.1 is characterized by a bright centrally condensed synchrotron nebula with relatively flat energy spectral index (∼2′ diameter, $\alpha = +0.12$ with $F_\nu \propto \nu^{-\alpha}$) and a radio shell with a steeper radio spectrum (∼8′ diameter; $\alpha = +0.77$). The synchrotron core is powered by the loss of rotational energy from a central pulsar; hereafter we refer to this core as the Pulsar Wind Nebula (PWN) in G0.9+0.1.

Owing to the high interstellar absorption ($\mathcal{N}_{\rm H} \sim 10^{23}$ cm^{-2}), G0.9+0.1 was barely detected by the *Einstein* Observatory (Helfand & Becker 1987). However, a much more convincing detection was obtained at higher energies with the *BeppoSAX* satellite (Mereghetti et al. 1998). Later on, Sidoli et al. (2000) interpreted this hard X-ray emission as non-thermal in origin. The small angular extent of the X-ray source (radius ∼1′), combined with an estimated age of the remnant of a few thousand years, is further evidence that the central radio core is powered by a young pulsar (∼2700 yr; Sidoli et al. 2000). The SNR shell remains unobserved in X-rays, probably due to the suppression of its soft X-ray flux by the line-of-sight absorption.

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Recently, G0.9+0.1 was observed with *Chandra* (Gaensler et al. 2001). The high angular resolution of *Chandra* provided an unprecedented view of the X-ray morphology of this SNR and identified a semi-circular arc-like feature, a jet-like feature and a very bright and small central emission region close to a very faint unresolved source (CXOU J174722.8-280915). The latter is inferred to be the pulsar itself (Gaensler et al. 2001). This X-ray morphology is similar to other PWNe powered by young pulsars, which display jets combined with a torus structure (Crab: Brinkmann et al. 1985; Weisskopf et al. 2000; Vela: Pavlov et al. 2000). As mentioned by Gaensler et al. (2001), the limited counting statistics of the 35 ks *Chandra* observation did not allow them to carry out a detailed spectral study of the various structures of the PWN, as ideally is required for the testing of current models.

Here we present the results from an observation of the composite SNR G0.9+0.1 made with *XMM-Newton*. We focus on the spectral analysis of different regions of the remnant on both relatively large and small spatial scales.

2. Observations and data analysis

G0.9+0.1 was observed by *XMM-Newton* on-axis on September 23–24, 2000 (∼17.2 ks and ∼12 ks for the MOS and PN cameras, respectively). The EPIC-MOS cameras were operated in the standard full-frame mode and the EPIC-PN

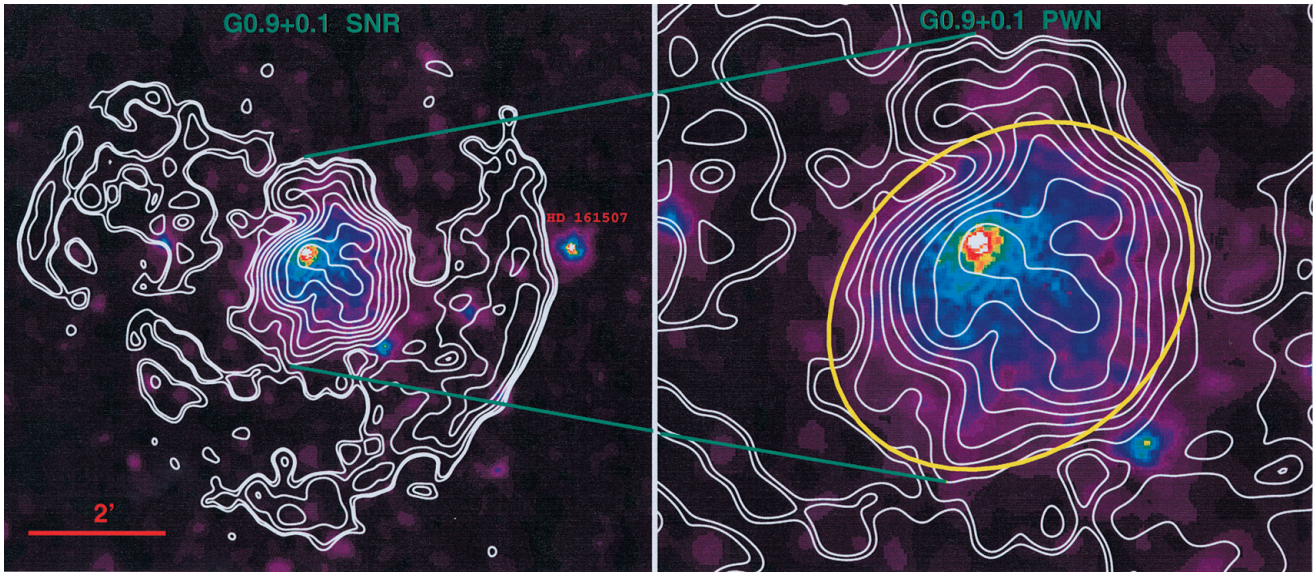

Fig. 1. *XMM-Newton* EPIC image of SNR G0.9+0.1 and its PWN in the energy band 1.5–12 keV with an adaptative smoothing filter with a signal-to-noise requirement of 5 and Gaussian smooth of 10″. The *VLA* radio contours at 1.5 GHz (20 cm) are superimposed in white. *Left panel*: overall remnant. *Right panel*: Pulsar Wind Nebula (PWN). The yellow ellipse represents the region taken for spectra analysis of the overall PWN.

camera in the extended full frame mode. Each camera was operated with the medium filter deployed.

Using the XMM Science Analysis Software (SAS version 5.2), the recorded events were screened by rejecting the high background periods which occasionally arise due to an intense incident flux of soft protons. After this data cleaning, the useful observing times are respectively for MOS1 and MOS2 about 15.5 ks and 15.4 ks, and 10 ks for PN. X-ray events corresponding to pattern 0–12 for the two MOS cameras were used, whereas for the PN only pattern 0–4 (single and double pixel events) were accepted. The astrometry of this observation was substantially improved using as reference the bright foreground F3V star HD 161507, which is an eclipsing binary of Algol type, detected at the western side of G0.9+0.1 (see Fig. 1, left panel).

For the imaging analysis we have summed up the MOS and PN data. We have used for the adaptative smoothing the tool asmooth from the SAS.

Our spectral analysis combines fits of the MOS and PN data (except in Sect. 3.3) employing the following camera response matrices: m1_medv9q20t5r6_all_15.rsp, m2_medv9q20t5r6_all_15.rsp, epn_ef20_sdY9_medium.rsp. We subtract from the source and the local background in our pointing, complementary data from a blank-field observation (kindly provided by David Lumb) in order to take into account the particle background (Majerowicz et al. 2002). The normalization between our pointing and the blank-field files is determined by their count ratios at (10–12) keV for the MOS and at (12–14) keV for the PN. For the vignetting, we applied the weighting method described by Arnaud et al. (2001). Our spectra were binned to 3σ before background subtraction, except in Sect. 3.1. xspec v11.1.0 was used for the spectral analysis. All subsequent errors are quoted at 90% confidence. Abundances are those of Anders & Grevesse (1989).

3. The composite SNR

The radio data of SNR G0.9+0.1 clearly show a radio shell and a very bright centrally condensed synchrotron nebula (Helfand & Becker 1987). Within the PWN, two central (west and east) peaks are present at 20 cm, while at 6 cm only the west peak is visible (Figs. 1a, b in Helfand & Becker 1987). Up to now, there is no known X-ray counterpart to the radio shell, while the PWN is clearly detected above about 2–3 keV (Mereghetti et al. 1998; Gaensler et al. 2001).

Figure 1 shows the *XMM-Newton* EPIC images of the composite SNR G0.9+0.1 in the energy band 1.5–12 keV with the radio contours overlaid. We use the radio data initially from Helfand & Becker (1987) which have been re-analyzed by Gaensler et al. (2001). The spatial resolutions are respectively 8″ for *XMM-Newton* and 40″ for the *VLA* data. Several X-ray bright point sources are observed within the field such as HD 161507 (a F3V star) and a bright X-ray source located to the south-west of G0.9+0.1 which appears to correspond to the 2MASS source 2MASS1747178-281025 (according to the cross-correlation results from *XMM-Newton* EPIC pipeline processing). A more focused spectral study of the radio shell region brings for the first time an indication of diffuse X-ray emission, which will be addressed in more detail below. Furthermore, the large scale X-ray morphology of the PWN ($R \sim 1'$) is in good agreement with the *VLA* radio contours at 20 cm (Fig. 1, right panel).

3.1. The SNR shell

We extracted spectra of the region corresponding to the radio shell ($R \sim 4'$), excluding the PWN (see ellipse in Fig. 1, right panel). We adopt a 3σ binning above background. Since the spectrum is strongly absorbed below 2 keV, due to the

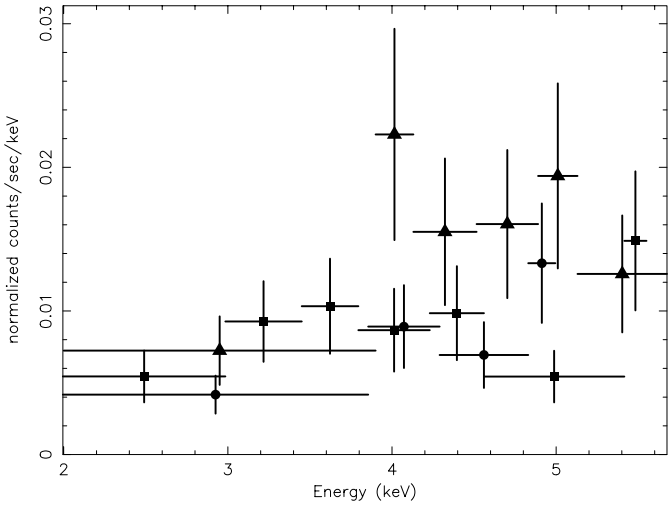

Fig. 2. XMM-Newton spectra (MOS1: filled circles, MOS2: filled squares, and PN: filled triangles) of the shell region ($R \sim 4'$, excluding the PWN) of G0.9+0.1.

interstellar medium in the Galactic Center region, we focus on the energy band 2–10 keV. The background taken for the shell analysis is an annulus centered on the PWN within radii 4.4′ and 8.8′, excluding the bright sources. The background subtracted spectra shown in Fig. 2, clearly reveal a signal above zero level which can be interpreted as the first detection of X-ray emission in this region. We fitted the combined MOS and PN spectra with both a non-thermal model (POWER-LAW) and a thermal model (MEKAL) with $\mathcal{N}_{\rm H}$ fixed at $1.39 \times 10^{23}\,{\rm cm}^{-2}$ as determined for the PWN (see Sect. 3.2). We find for the MEKAL model $kT = 3.4^{+5.3}_{-1.2}$ keV (χ^2/d.o.f. = 17.3/15), and for the power-law model $\Gamma = 2.6^{+0.7}_{-0.8}$ (χ^2/d.o.f. = 17.4/15). The inferred temperature is relatively high for a several thousand year old SNR. Both models give respectively an unabsorbed flux (2–10 keV), of $2.2^{+4.1}_{-1.5} \times 10^{-12}$ erg cm^{-2} s^{-1}, and $2.3^{+1.4}_{-0.9} \times 10^{-12}$ erg cm^{-2} s^{-1}. The inferred fluxes are in good agreement with the upper limit quoted by Sidoli et al. (2000), i.e. 3×10^{-12} erg cm^{-2} s^{-1}. We also carried out a spectral analysis (fixing $\mathcal{N}_{\rm H}$ at $1.39 \times 10^{23}\,{\rm cm}^{-2}$) of the brightest western region of the radio shell (the statistic of the eastern part is too poor to allow a spectral analysis). Again, this region is dominated by the local diffuse emission of the GC region and thus requires appropriate background subtraction, which we extract from an adjacent field. The corresponding fit results for the temperature are $kT = 2.7^{+17.6}_{-1.2}$ keV (χ^2/d.o.f. = 3.5/5), and is thus not well constrained. A power law model yields $\Gamma = 2.9^{+1.2}_{-1.3}$ (χ^2/d.o.f. = 3.3/5).

We extracted an image in the energy band 3.9–5.8 keV, i.e. in the spectral interval with maximum emission. However, no distinct borderline is detected between the X-ray shell and the diffuse medium located in the GC region.

3.2. The PWN

We observe that the core of the PWN is surrounded by extended diffuse emission (Fig. 1, right panel). The position of the X-ray core matches the eastern radio peak while there is no bright X-ray counterpart to the western radio peak. Figure 3 displays the PWN in different energy bands, namely 1.5–3 keV, 3–6 keV, and 6–12 keV. In the energy band 1.5–3 keV, there is little evidence for a dominant core and the radio and X-ray morphologies are similar (see Fig. 3, left panel). The core of the PWN is only observable above ∼3 keV (Fig 3, middle and right panels). At high energy (6–12 keV) the symmetry axis is analogous to the one of *Chandra* data, the western radio lobe is very toned down, and the faint X-ray source CXOU J174722.8-280915 is apparent (Fig. 3, right panel).

Table 1. Results of the spectral fits for the X-ray PWN. The assumed interstellar photo-electric absorption cross-sections are from Morrison & McCammon (1983; WABS in XSPEC) or from Wilms et al. (2000; TBABS). Uncertainties are quoted at 90% confidence. The unabsorbed fluxes (2–10 keV) are expressed in 10^{-12} erg cm^{-2} s^{-1}.

Model	$\mathcal{N}_{\rm H}$ (10^{23} cm^{-2})	Γ or kT (keV)	χ^2 (446 d.o.f.)	$F^{\rm unabs}_{2-10}$
wabs*pow	$1.47^{+0.14}_{-0.13}$	$1.99^{+0.19}_{-0.18}$	379.6	5.78
wabs*brems	$1.33^{+0.11}_{-0.10}$	$10.7^{+4.2}_{-2.3}$	379.1	5.16
tbabs*pow	$1.39^{+0.13}_{-0.12}$	$1.93^{+0.18}_{-0.18}$	378.5	5.73
tbabs*brems	$1.27^{+0.10}_{-0.10}$	$11.7^{+4.9}_{-2.7}$	378.0	5.17

For comparison with previous studies, we fit the spectrum of the overall PWN. The region considered is defined by the ellipse displayed in Fig. 1 (right panel). Our spectra were binned to 3σ before background subtraction. We defined a local background corresponding to the central MOS 1 CCD region, excluding the bright point sources and the central part within a radius of 2.6′. As shown in Fig. 4, the corresponding spectra extend up to almost 12 keV for the PN data, and to about 9 keV in the MOS data, allowing strong constraints to be placed on the slope of the continuum. The data are well fitted by either an absorbed power-law or a thermal bremsstrahlung model. For the photo-electric absorption ($\mathcal{N}_{\rm H}$), we use first the cross-sections of Morrison & McCammon (1983, WABS). The best-fit parameters are reported in Table 1. We also applied the same spectral fitting to spectra with larger binsizes (higher signal to noise) and obtained fit results, which were in very good agreement with the values quoted here. The best-fit parameters relating to the power-law are compatible, within the error bars, with the *BeppoSAX* and *Chandra* values (Sidoli et al. 2000: $\mathcal{N}_{\rm H} = 1.09^{+0.24}_{-0.21} \times 10^{23}\,{\rm cm}^{-2}$, $\Gamma = 1.95^{+0.33}_{-0.30}$; Gaensler et al. 2001: $\mathcal{N}_{\rm H} = 1.6 \pm 0.02 \times 10^{23}\,{\rm cm}^{-2}$, $\Gamma = 2.4 \pm 0.4$), though more tightly constrained.

In the following, we use the updated cross-sections for X-ray absorption by the interstellar medium (TBABS in XSPEC) from Wilms et al. (2000). The best-fit parameters although in good agreement with the previous ones (see Table 1), give systematically slightly lower $\mathcal{N}_{\rm H}$ values using TBABS absorption model than WABS absorption model. We obtain for the power-law model a luminosity in the 2–10 keV energy band of $6.5 \times 10^{34}\,d_{10}^2$ erg s^{-1} (d_{10} is the distance in units of 10 kpc).

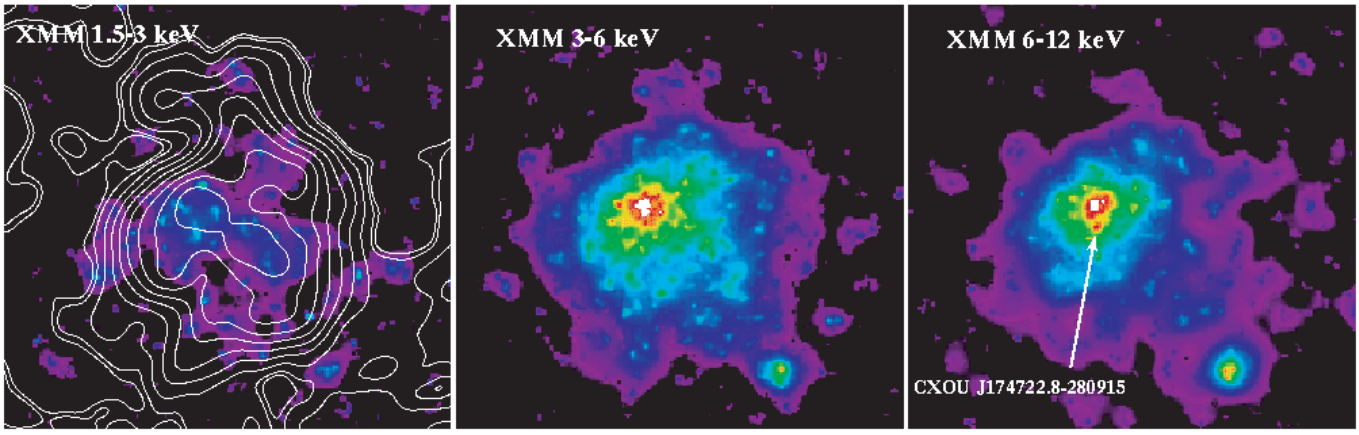

Fig. 3. *XMM-Newton* EPIC image of the PWN in different energy bands, obtained with adaptative smoothing with a signal-to-noise requirement of 5 and a Gaussian smooth of 20″. The *VLA* radio contours at 1.5 GHz (20 cm) are superimposed in white.

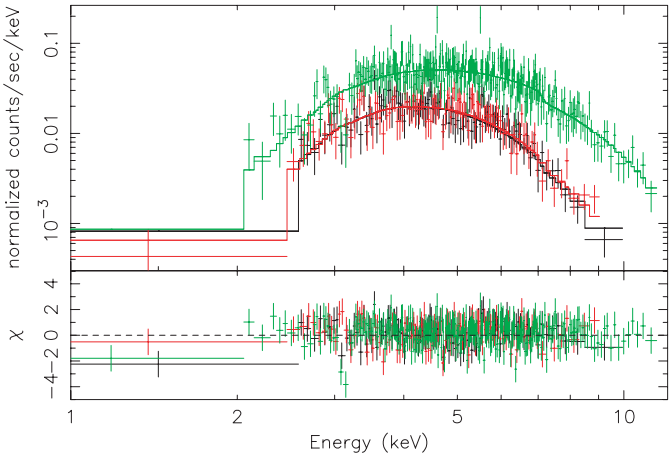

Fig. 4. Combined fit (MOS1 in black, MOS2 in red, and PN in green) of the PWN of G0.9+0.1 with an absorbed power law (TBABS*PO). See Table 1 for the inferred parameters.

3.3. CXOU J174722.8-280915

CXOU J174722.8-280915 is a point source located about 8″ below the X-ray bright core of the PWN in the *Chandra* observations and has been identified as the pulsar itself by Gaensler et al. (2001). In our *XMM-Newton* data, we are able to pick out this point source only above 6 keV (Fig 3, right panel).

We extracted a spectrum from a circular region ($R = 5''$) centred on the position of the point source (Fig 3, right panel). For the background we took a region near the source avoiding any strong contamination from the core emission (due to the point spread function of the EPIC cameras). We used only the spectra from the MOS cameras which benefit from the best mirrors and pixel sampling. As in Gaensler et al. (2001), we explore the possibility that CXOU J174722.8-280915 corresponds to a young central pulsar. Its X-ray emission is then expected to be due to either modified black-body emission from the neutron star surface or non-thermal emission from the neutron star magnetosphere (e.g. Becker & Trümper 1997). As calculated by van Ripert et al. (1995), a neutron star with an age between 10^3 and 10^4 years should have a surface temperature of $kT \sim 0.1$ keV. Fixing $N_{\rm H}$ at 1.39×10^{23} cm^{-2} and the temperature at 0.1 keV, we obtain a very bad fit with $\chi^2_{\rm red} = 2.4$ (5 d.o.f.). Letting kT or $N_{\rm H}$ free does not improve the fit. In contrast when we assume a non-thermal model (as in Gaensler et al. 2001), with a photon index of $\Gamma = 1.5$ (typical for magnetospheric emission, Becker & Trümper 1997) and $N_{\rm H} = 1.39 \times 10^{23}$ cm^{-2}, we obtain a much better fit with $\chi^2_{\rm red} = 0.8$ (5 d.o.f.). As for the black-body model, letting Γ or $N_{\rm H}$ free does not lead to an improvement in the fits. The corresponding absorption corrected flux in the 2–10 keV range is about 9.2×10^{-14} erg s^{-1}, implying a luminosity of about $10^{33}\, d^2_{10}$ erg s^{-1}.

4. Large-scale structure in the PWN: Tests of the plerionic interpretation

On large-scales the X-ray morphology of the PWN determined from the *XMM-Newton* observations is in relatively good agreement with the overall radio emission. The *XMM-Newton* data enable us for the first time to study the variation of the spectral index within the PWN. In other plerions, a softening of the spectrum with increasing radius has been observed; for example in 3C 58 (Torii et al. 2000; Bocchino et al. 2001), G21.5-0.9 (Slane et al. 2000; Warwick et al. 2001) and IC 443 (Bocchino & Bykov 2001). This spectral softening can be explained by the shorter lifetime of high energy electrons compared to lower energy electrons. To look for this effect, we extract the spectra in four regions, which are displayed in Fig. 5 (left panel). The outer ellipse corresponds to the region taken for the spectral analysis of the overall PWN (see Sect. 3.2). The observed spectra and best-fit models for regions 1 and 3 are shown, as examples, in Fig. 2 in Porquet et al. (2002). We fit the spectra of these regions with an absorbed power-law model (TBABS*PO), fixing $N_{\rm H}$ to the value obtained for the overall PWN, i.e. $N_{\rm H} \sim 1.4{\times}10^{23}$ cm^{-2} (see Table 1). The corresponding best-fit parameters are reported in Table 2. The value of the photon index Γ shows a clear steepening of the spectrum from the inner part toward the outer part of the PWN (Fig. 5, right panel). Freeing the absorption column density does not result in significant better fits. There is no evidence for any

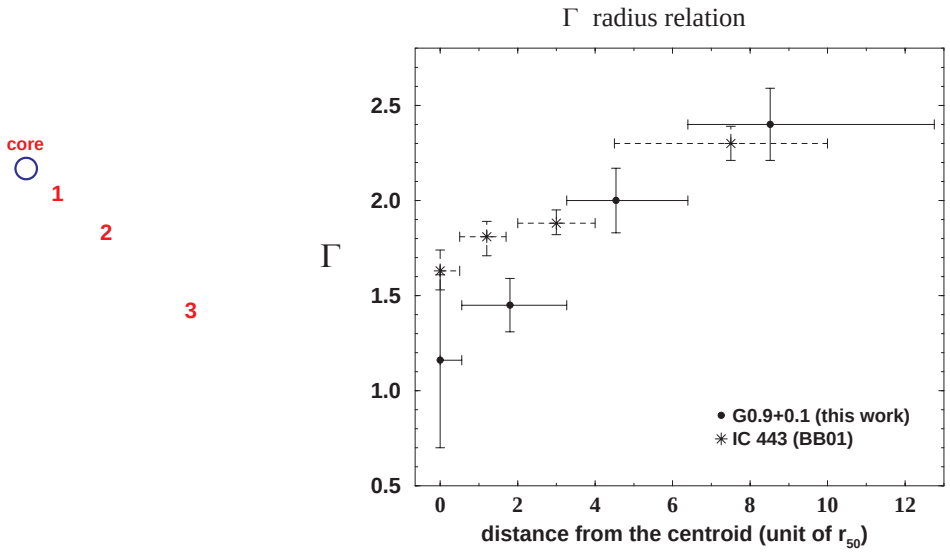

Fig. 5. *Left panel*: *XMM-Newton* EPIC image ($2.7' \times 2.7'$) of the PWN within G0.9+0.1, in the 3–8 keV energy range, obtained with an adaptive smoothing filter with a signal-to-noise requirement of 5, and Gaussian smoothing of $10''$. The regions used for the spectral analysis are superposed: core, regions 1, 2, and 3. *Right panel*: Photon index versus distance from the core. The *X*-axis (defined in Bocchino & Bykov 2001) shows the weighted mean distance of the pixels of a given region from the centroid of the nebula, expressed in units of r_{50}, the radius at which the plerion surface brightness drops by a factor of 2 ($7.15''$ for G0.9+0.1). The r_{50} unit gives a measure independent of the distance to the nebulae. For comparison, data for IC 443 are also plotted (taken from Bocchino & Bykov 2001).

Table 2. Best-fit parameters for the combined (MOS+PN) spectrum of the large-scale structures of the PWN regions (defined in Fig. 5, left panel). The model is an absorbed power law model (TBABS*PO). The cross-sections of the interstellar absorption are from Wilms et al. (2000). Unabsorbed X-ray fluxes (2–10 keV) are expressed in 10^{-12} ergs cm^{-2} s^{-1}. *At the top*: $\mathcal{N}_{\rm H}$ is frozen to the value obtained for the entire PWN (see Table 1). *At the bottom*: $\mathcal{N}_{\rm H}$ is a free parameter.

	$\mathcal{N}_{\rm H}$ (10^{23} cm^{-2})	Γ	χ^2/d.o.f.	$F_{2-10}^{\rm unabs}$
core	1.39	$1.16^{+0.45}_{-0.46}$	18.2/18	0.24
region 1	1.39	$1.45^{+0.14}_{-0.14}$	105.7/146	1.80
region 2	1.39	$2.00^{+0.17}_{-0.17}$	121.2/136	1.66
region 3	1.39	$2.40^{+0.19}_{-0.19}$	162.2/173	2.02
core	$1.36^{+0.82}_{-0.61}$	$1.12^{+1.11}_{-0.88}$	18.2/17	0.24
region 1	$1.54^{+0.25}_{-0.22}$	$1.61^{+0.30}_{-0.28}$	104.5/145	1.95
region 2	$1.53^{+0.27}_{-0.24}$	$2.18^{+0.37}_{-0.35}$	120.2/135	1.83
region 3	$1.35^{+0.24}_{-0.22}$	$2.35^{+0.39}_{-0.35}$	162.1/172	1.96

absorption variation across the PWN, as the column densities are compatible to within 12% for the different regions of the nebula.

The spectral softening observed from the core to the outskirts of the PWN is consistent with synchrotron radiation losses of high energy electrons as they diffuse through the nebula. For comparison we show in Fig. 5 (right panel) the radial variation of the spectral index measured for IC 443 (Bocchino & Bykov 2001). Our data show indication of a harder core and a stronger softening from the core towards the outer part of the G0.9+0.1 nebula, than is seen in IC 443, and also in other PWNe such as 3C 58 and G21.5-0.9 (see Fig. 5 in Bocchino & Bykov 2001). This suggests the possibility of a relatively strong magnetic field existing in the G0.9+0.1 pulsar and/or its SNR environment. The variation in the spectral index between the core and the PWN is roughly consistent with the empirical linear relationship derived by Gotthelf & Olbert (2002) for 6 pulsars and their respective PWNe, i.e. $\Gamma_{\rm PWN} = 0.8 \times \Gamma_{\rm core} + 0.8$. SNR G54.1+0.3 shows a similar spectral index, $\Gamma = 1.09^{+0.08}_{-0.09}$, for its pulsar region (Lu et al. 2002), and $\Gamma = 1.9 \pm 0.2$ for the overall PWN (Lu et al. 2001).

5. Small-scale structures in the PWN

The high angular resolution of *Chandra* has provided an unprecedented view of the X-ray morphology of this PWN. Relatively small-scale features identified in the *Chandra* observations (Gaensler et al. 2001) include a semi-circular arc-like feature, a jet-like feature, and a bright clump close to a very faint unresolved source (CXOU J174722.8-280915, which as noted earlier is inferred to be the pulsar itself). This X-ray morphology is similar to other PWNe powered by young pulsars which display jets combined with a torus structure (e.g. Crab: Brinkmann et al. 1985; Weisskopf et al. 2000; Vela: Pavlov et al. 2000). The X-ray brightest region matches the eastern radio peak while the western side of the X-ray arc-like feature (structure 6 in Fig. 6) corresponds to the western radio peak. In the radio band, as displayed in Fig. 1 in Helfand & Becker (1987), there is a spectral index change between 20 cm and 6 cm, i.e. the eastern peak present at 20 cm disappears at 6 cm. The jet-like feature (structure 4 in Fig. 6) corresponds to a distortion in the radio contours.

In our *XMM-Newton* close-up view, these small-scale features are not as clearly detected as in the *Chandra* data.

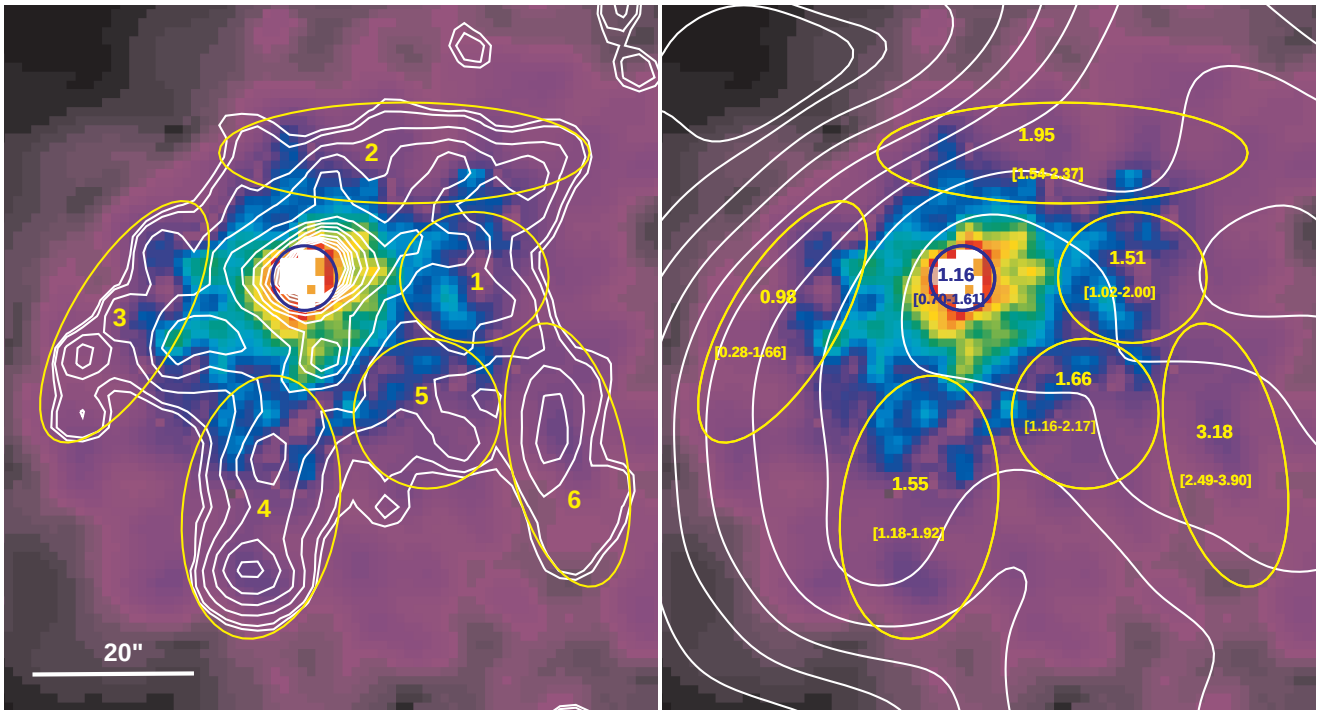

Fig. 6. Close-up view with XMM-Newton (EPIC) of the center of the PWN in the 3–8 keV energy band obtained with an adaptative smoothing filter with a signal-to-noise requirement of 5, and Gaussian smoothing of 10′. *Left panel*: white contours correspond to the Chandra observation in the same energy range. Yellow regions represent different small-scale structures surrounding the bright central core (dark blue circle) matching the Chandra contours (see Gaensler et al. 2001): "East arc-like feature" (structure number 3), "jet-like feature" (structure 4), and "West arc-like feature" (structure 6). *Right panel*: the radio *VLA* contours at 20 cm overlaid on the same XMM-Newton image. The corresponding spectral indices and the minimum and maximum values (in brackets) are given within 90% confidence uncertainties.

Table 3. Best-fit parameters for the combined (MOS+PN) spectrum of the small-scale structures of the PWN of G0.9+0.1 (defined in Fig. 6). The model is an absorbed power law model (TBABS*PO). Unabsorbed X-ray fluxes (2–10 keV) are expressed in 10^{-12} ergs cm^{-2} s^{-1}. $\mathcal{N}_{\rm H}$ is fixed at 1.39×10^{23} cm^{-2}.

structure	Γ	χ^2/d.o.f.	$F_{\rm X}^{(b)}$ (2–10 keV)	*Chandra* identification
core	$1.16^{+0.45}_{-0.46}$	18.2/18	0.24	
1	$1.51^{+0.49}_{-0.49}$	13.5/16	0.22	
2	$1.95^{+0.42}_{-0.41}$	21.8/25	0.32	
3	$0.98^{+0.68}_{-0.70}$	15.9/14	0.20	"East arc"?
4	$1.77^{+0.41}_{-0.41}$	13.6/21	0.28	"jet" ?
5	$1.66^{+0.51}_{-0.50}$	19.1/18	0.23	
6	$3.18^{+0.72}_{-0.69}$	13.8/12	0.22	"West arc"?

However, the extent of this emission is about a factor 2 larger than observed with *Chandra*, thanks to the high sensitivity of *XMM-Newton*. Using as guide the locations defined by *Chandra*, we are able for the first time to derive the spectra of the various small-scale structures seen in the G0.9+0.1 PWN. Determining the spectral variations in these small-scale structures can give an indication of the geometry and orientation of the nebula. The best-fit parameters of the combined fits (MOS and PN) of the six regions identified in Fig. 6 are given in Table 3. Figure 6 (right panel) also displays the spectral index measured for each region. The region corresponding to the eastern part of the arc-like feature, which was first revealed by *Chandra*, shows apart from the core region ($\Gamma = 1.2 \pm 0.5$) the hardest spectrum with a spectral index of ($\Gamma = 1.0 \pm 0.7$) among all the structures of the PWN. The former region encloses some unresolved bright knots which may be point-like sources, however if we exclude these bright sources we still obtain a very flat power-law form with $\Gamma \sim 1.1 \pm 0.8$ ($\chi^2_{\rm red} = 1.2$, 12 d.o.f). In contrast the region corresponding to the western side of the arc-like feature, appears to have the steepest photon index ($\Gamma \sim 3.2 \pm 0.7$) of any region within the PWN. In order to investigate spectral variations of the small scale structures, we constructed a subtracted background hardness ratio image ([6–10] keV/[3–6] keV) of the central region (see Fig. 7). It can be clearly seen that the eastern part of the arc-like feature shows harder X-ray emission than the western part. A possible explanation is that the east arc-like feature is pointing towards the observer, and that its spectral hardness is due to the relativistic beaming or Doppler boosting of the electrons, while the western arc corresponds to the opposite part of the torus suggested in the *Chandra* data. The region associated with the jet-like feature in Gaensler et al. (2001) does not exhibit a harder spectrum as suggested by their inferred hardness ratio. Splitting this structure into two parts, north and south, and fitting the spectra separately gives respectively as spectral indexes $\Gamma \sim 1.4 \pm 0.9$, and $\Gamma \sim 2.1 \pm 0.5$. The possible observed curvature of this putative jet from small-scale to large-scale may be the result of an

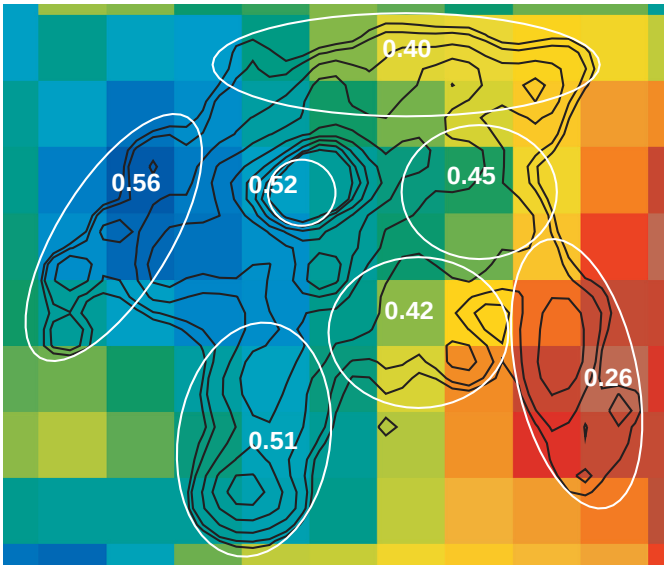

Fig. 7. Background subtracted hardness ratio map ([6–10] keV/[3–6] keV) of the small-scale structures with a pixel size of 8″. The colour coding goes from blue (hard spectrum) to red (soft spectrum). To guide the eye we give explicit hardness ratio values for characteristic pixels. As in Fig. 6 (*left panel*), the Chandra contours and the regions taken for our spectral analysis are superposed.

interaction with the high magnetic fields of the Galactic Center region.

6. Summary

We present first results of an *XMM-Newton* observation of the composite supernova remnant G0.9+0.1 located in the Galactic Center region.

The high sensitivity of the *XMM-Newton* observatory allows for the first time to detect diffuse X-ray emission in the region of the radio shell. We find an unabsorbed flux of about 2×10^{-12} erg cm^{-2} s^{-1} (2–10 keV).

The *XMM-Newton* X-ray morphology is in relatively good agreement with published (*VLA*) radio maps of the PWN. We have carried out the first detailed X-ray spectral analysis of the PWN inside G0.9+0.1. On large scales, there is a clear softening of the spectrum with radial distance from the core, a phenomenon which has been observed previously in a number of X-ray plerions (e.g., 3C58, G21.5-0.9, IC443). Detailed spectral analysis is also presented of the small-scale structures evident in the PWN. The eastern part of the arc-like feature (revealed by *Chandra*) presents a reliable indication for a very hard photon index ($\Gamma \sim 1.0 \pm 0.7$), while the western part presents a very soft spectrum ($\Gamma \sim 3.2 \pm 0.7$). A possible explanation is that the east arc-like feature is pointing towards the observer, and that its spectral hardness is due to the relativistic beaming or Doppler boosting of the electrons, while the western arc corresponds to the opposite part of the torus suggested in the *Chandra* data.

G0.9+0.1 provides further clues concerning the processes by which pulsars connect with their environment and is also important to illustrate a possible impact of the strong external magnetic field that pervades the Galactic Center region.

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