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Molecules in the Cep E-mm jet: evidence for shock-driven photochemistry ?

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ABSTRACT

The chemical composition of protostellar jets and its origin are still badly understood. More observational constraints are needed to make progress. With that objective, we have carried out a systematic search for molecular species in the jet of Cep E-mm, a template for intermediate-mass Class 0 protostars, associated with a luminous, high-velocity outflow. We made use of an unbiased spectral line survey in the range 72–350 GHz obtained with the IRAM 30m telescope, complementary observations of the CO $J=3-2$ transition with the JCMT, and observations at $1''$ angular resolution of the CO $J=2-1$ transition with the IRAM Plateau de Bure interferometer. In addition to CO, we have detected rotational transitions from SiO, SO, H₂CO, CS, HCO⁺ and HCN. A strong chemical differentiation is observed in the southern and northern lobes of the jet. Radiative transfer analysis in the Large Velocity Gradient approximation yields typical molecular abundances of the order of 10^{-8} for all molecular species other than CO. Overall, the jets exhibit an unusual chemical composition, as CS, SO and H₂CO are found to be the most abundant species, with a typical abundance of $(3-4) \times 10^{-8}$. The transverse size of the CO jet emission estimated from interferometric observations is about 1000 au, suggesting that we are detecting emission from a turbulent layer of gas entrained by the jet in its propagation and not the jet itself. We propose that some molecular species could be the signatures of the specific photochemistry driven by the UV radiation field generated in the turbulent envelope.

Key words: physical data and processes: astrochemistry – ISM: jets and outflows – molecules-abundances – Stars:formation

1 INTRODUCTION

Protostellar outflows are one of the most spectacular manifestations of the mass-loss phenomena which take place along the protostellar evolution, from the early, embedded Class 0 to the late Class I phase. Many protostellar outflows show evidence for molecular jets. The origin of molecules in protostellar jets is strongly debated and observational constraints are lacking. This issue is all the more important as the jet molecular composition could bring useful constraints on the gas dynamics and possibly on the launching mechanism itself. Numerical simulations suggest that turbulent entrainment from the surrounding envelope is not efficient enough to explain the CO column densities (Taylor et al. 1995). Several other mechanisms have been proposed: formation by ion chemistry at the base of an atomic stellar wind (Glassgold et al. 1991), formation behind atomic jet

shocks (Raga et al. 2005), direct MHD ejection of molecules from the disk (Panoglou et al. 2012). Therefore, a census of the molecular composition of protostellar jets could help discriminate between these various scenarios.

The Large Program CALYPSO has investigated the emission of a few molecular tracers between 92 and 232 GHz with the IRAM Plateau de Bure interferometer towards a sample of 16 nearby ($d \leq 300$ pc) Class 0 protostars objects (Maury et al. 2015). It comes out that jets are often associated with SiO and SO, in addition to CO (see Codella et al. 2014; Santangelo et al. 2015). Until now, there have been very few attempts to investigate the molecular content of these jets. Tafalla & Hacar (2010) have led a molecular survey of the content of the two protostellar outflows L1448-mm and IRAS 04166+2706, from which they reported the detection of SiO, SO, HCN, and CH₃OH. The molecular composition differs markedly between the low-velocity bipolar outflow and the protostellar jet itself. This point is well illustrated by the L1157 low-velocity outflow which was long

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identified as "chemically active" (Bachiller & Perez-Gutierrez 1997; Codella et al. 2010; Lefloch et al. 2017) whereas the jet molecular emission was detected only recently, in the lines of SiO and CO (Tafalla et al. 2015; Podio et al. 2016).

In this context, we have conducted an unbiased molecular line survey of the high-velocity jet and outflow of Cep E-mm, a template for intermediate-mass Class 0 protostar. The source has a luminosity of about $100 L_{\odot}$ (Lefloch et al. 1996; Chini et al. 2001) and is located at a distance of 730 pc in the Cepheus-E molecular cloud (Sargent 1977). Both the source and its outflow have been extensively observed in the millimetre to far-infrared wavelengths range (see, e.g., Eis-löffel et al. 1996; Lefloch et al. 1996; Noriega-Crespo et al. 1998; Hatchell et al. 1999b; Moro-Martin et al. 2001; Smith et al. 2003; Lefloch et al. 2015; Ospina-Zamudio et al. 2018). However the molecular content of Cep E-mm have received only little attention until the systematic study by Ospina-Zamudio et al. (2018). These authors have reported the detection of a hot corino in the source. They have shown that Cep E-mm is actually a binary system, whose components Cep E-A and Cep E-B both drive high-velocity molecular jets. The authors identified Cep E-A as the powering source of the high-velocity jet associated with HH377. They also found that the outflow is chemically active, following the empirical classification proposed by Bachiller et al. (1997). The physical properties of the southern lobe of the outflow, both the entrained gas and the jet, were obtained from a multi-line analysis between $J=1-0$ and $J=16-15$, of the CO jet emission (Lefloch et al. 2015). This study revealed the presence of hot, dense CO knots along the jet, which were interpreted as the signature of internal shocks.

In this article, we report the results of our unbiased search for molecular emission in the high-velocity jet of Cep E-A. The article is organized as follows. In Section 2, we present the observations carried out with the IRAM 30m telescope, the JCMT and the IRAM Plateau de Bure Interferometer (PdBI). In Section 3, we present the results on the jet molecular content. In addition to the previously known molecules in the jet (CO, SiO), we report on the detection of SO, CS, H_2CO , HCO^+ and HCN. In Section 4, we derive the physical conditions and the molecular abundances of the various tracers detected in the jet. In each outflow lobe, the jet composition is compared with that of the low-velocity outflow cavity. We discuss the origin of the jet molecular emission. The specific jet molecular composition is found to agree qualitatively with the predictions of a simple shock-driven UV photochemical model. Our conclusions are summarized in Section 5.

2 OBSERVATIONS

2.1 IRAM 30m

The observational data towards the protostar on coordinates $\alpha(2000) = 23^h03^m12^s.8$, $\delta(2000) = 61^\circ42'26''$ have already been presented in Ospina-Zamudio et al. (2018). As recalled by the authors, calibration uncertainties in the 3, 2, 1.3 and 0.9mm bands are typically 10, 15, 20 and 30%, respectively. In this work, line intensities are expressed in units of antenna temperature corrected for the atmospheric absorption T_A^* . The rms noise per velocity interval of 1 km s^{-1}

expressed in units of T_A^* lies in the range 2–5 mK in the 3mm band, 3–7 mK in the 2mm and 1.3mm band, and 15–20 mK in the 0.9mm band.

The observational parameters and the parameters of all the detected lines are given in Table 1. Molecular line brightness temperatures T_B were estimated from main-beam temperatures T_{MB} , following the procedure described in Lefloch et al. (2015) to estimate the filling factor of the molecular jet emission in the telescope main-beam. We adopted the main-beam efficiency values η_{mb} , as monitored by IRAM¹.

2.2 JCMT

The CO lines $J=3-2$ 345.796 GHz and $J=4-3$ 461.041 GHz were observed toward the protostar position at the JCMT in June 1997. The CO $J=4-3$ was also observed at the offset position ($0'', +12''$). The observational procedures and data reduction procedures have been discussed in Lefloch et al. (2015).

2.3 PdBI

The CO $J=2-1$ line at 230.538 GHz was observed with the five-antenna array of the PdBI in March 1997 in configurations C and D, and again in January and February 1998 in configurations A and B. The details of the observing procedure and data reduction have been presented in Lefloch et al. (2015). As a result of the calibration conducted under CLIC², we obtained a natural weighted map with a synthetic beam of $1.07'' \times 0.87''$ (PA= 73°).

3 THE HIGH-VELOCITY JET

Analysis of the CO $J=2-1$ interferometric observations (Lefloch et al. 2015) revealed three spatial components in the protostellar outflow: the molecular jet (Fig. 1), the outflow cavity (Fig. 2), and the terminal shock HH377. Each of these components appear to have a specific spectral signature.

The molecular jet (hereafter the high-velocity jet) powered by the protostar Cep E-A (Ospina-Zamudio et al. 2018) is detected between $v_{lsr} = -135$ and -110 km s^{-1} and between $+50$ and $+80 \text{ km s}^{-1}$ towards the southern and northern lobes, respectively (Fig. 1).

The outflow cavity is detected at velocities from ambient ($v_{lsr} = -10.9 \text{ km s}^{-1}$) up to -100 km s^{-1} ($+60 \text{ km s}^{-1}$) in the southern (northern) lobe. The bulk of the emission is detected up to 10 km s^{-1} from the ambient cloud velocity (hereafter the low-velocity outflow) and arises from the outflow cavity walls (Fig. 2).

A systematic search down to the 3σ level yielded detection of the following molecular species in the high-velocity jet: CO, SiO, SO, HCO^+ , HCN, CS and H_2CO (Fig. 3). All the molecular transitions detected in the jet are listed in Table 1 and a montage of the lowest excitation transitions of each molecular species is displayed in the right panel of

¹ <http://www.iram.fr>

² <http://www.iram.fr/IRAMFR/GILDAS/>

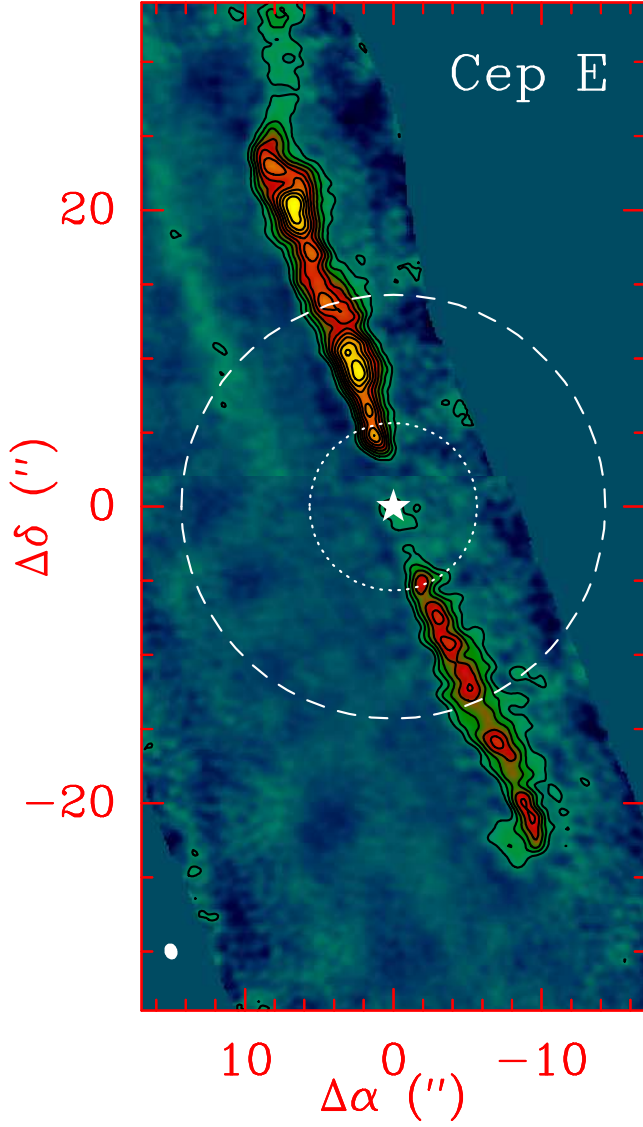


Figure 1. CO $J=2-1$ velocity-integrated emission of the protostellar jet observed with the PdBI at $1''$ resolution (from Lefloch et al. 2015). The emission is integrated between -135 and -110 km s^{-1} and between $+50$ and $+80$ km s^{-1} towards the southern and northern lobes, respectively. The first contour and contour interval are 20% and 10% of the peak intensity. The synthetic beam size (HPBW) is drawn by a filled ellipse. The IRAM 30m telescope beam size (HPBW) at the frequency of the SiO $J=2-1$ and $J=5-4$ is drawn by a dashed and dotted circle, respectively. A white star marks the position of the protostar Cep E-A.

Fig. 3. The spectra of all the transitions detected are displayed in Figs. A1–A4. Inspection of the line profiles in Fig. 3 and Figs. A1–A4 immediately show:

- different chemical properties between both jet lobes: For instance, H_2CO and CS lines are bright in the northern lobe, while these species are not detected in the southern lobe. On the contrary, HCO^+ is detected in both lobes.
- different physical structures: different molecular species display different line profiles inside the same lobe and between both lobes.

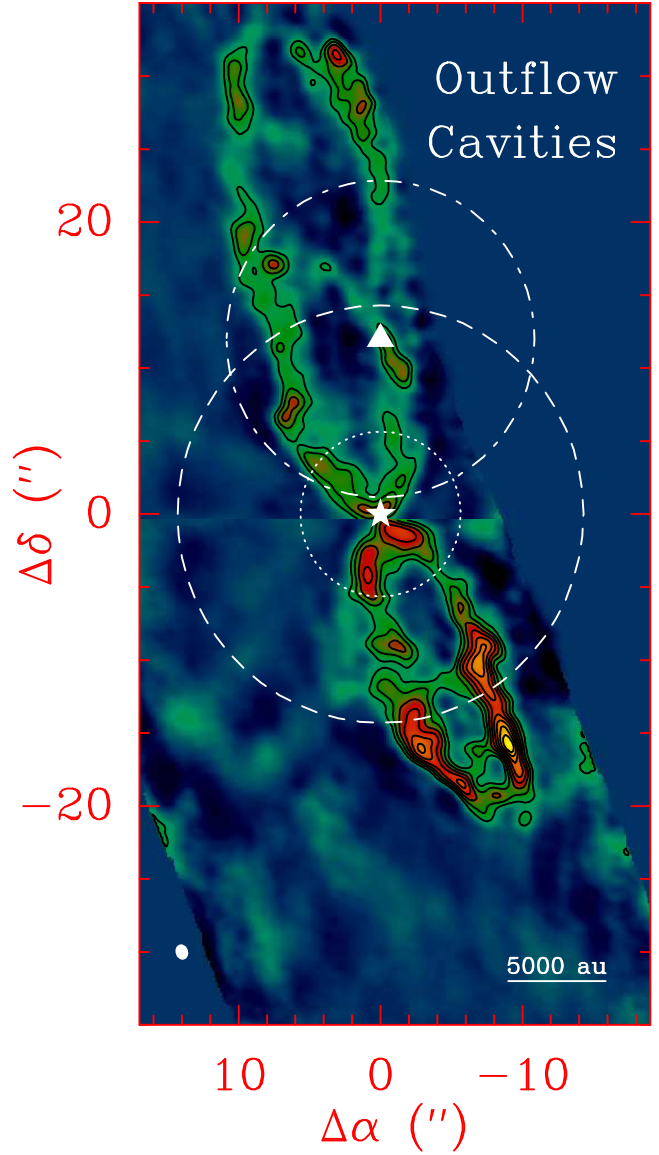


Figure 2. CO $J=2-1$ velocity-integrated emission of the low-velocity outflow observed with the PdBI at $1''$ resolution. The emission is integrated between -20 and -14 km s^{-1} and between -8 and -6 km s^{-1} towards the southern and northern lobes, respectively. The first contour and contour interval are 20% and 10% of the peak intensity. The synthetic beam size (HPBW) is drawn by a filled ellipse. The IRAM 30m telescope beam size (HPBW) at the frequency of the SiO $J=2-1$ and $J=5-4$ is drawn by a dashed and dotted circle, respectively. The filled triangle marks the offset position $(0'', +12'')$ observed with the IRAM 30m and the JCMT. The dashed-dotted circle draw the IRAM 30m beam at the frequency of the CO $J=1-0$ line. A white star marks the position of the protostar Cep E-A.

3.1 CO

The single-dish CO line observations towards the offset position $(0'', +12'')$ and $(0'', 0'')$ show that the CO emission ranges between -150 km s^{-1} and $+80$ km s^{-1} (Fig. 3). The complex structure of the outflowing gas is evidenced by the presence of several high-velocity components in both outflow lobes, and the various spectral signatures detected with different molecular gas tracers. Whereas the ambi-

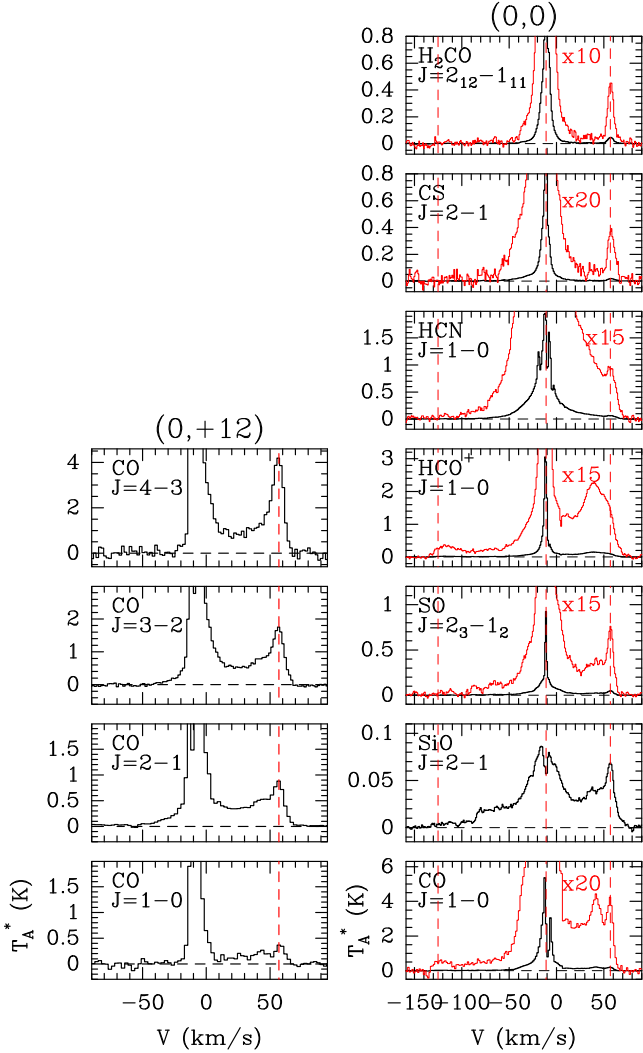


Figure 3. (*left*) CO line emission detected towards the outflow northern lobe at the offset position ($0'', +12''$). (*right*) Molecular line emission detected towards the protostar position ($0'', 0''$). Original spectra are displayed in black. Spectra with a magnifying factor are drawn in thin red lines. Fluxes are expressed in units of antenna temperature T_A^* , corrected for atmospheric absorption. The ambient cloud velocity ($v_{\text{lsr}} = -10.9 \text{ km s}^{-1}$) and the mean jet velocity in the northern and southern lobes ($v_{\text{lsr}} = +57 \text{ km s}^{-1}$ and -125 km s^{-1} , respectively) are marked by red, dashed lines.

ent cloud emission peaks at $v_{\text{lsr}} = -10.9 \text{ km s}^{-1}$, the emission of the protostellar jet lies in the velocity intervals $[+50; +80] \text{ km s}^{-1}$ and $[-150; -100] \text{ km s}^{-1}$ for the redshifted and blueshifted lobe, respectively and peaks at $+56 \text{ km s}^{-1}$ and -125 km s^{-1} , respectively (see also Lefloch et al. 1996, 2015). A map of the CO $J=2-1$ jet velocity-integrated emission between $[+50; +80] \text{ km s}^{-1}$ and $[-150; -100] \text{ km s}^{-1}$, respectively, obtained with the PdBI at $1''$ resolution is displayed in Fig. 1. The overall emission of both jet lobes is rather similar, with a clumpy structure suggestive of internal shocks and multiple ejections (see also Lefloch et al. 2015). We note that the CO $J=2-1$ line emission brightness in the northern lobe is larger than in the southern lobe by typically a factor of ~ 5 .

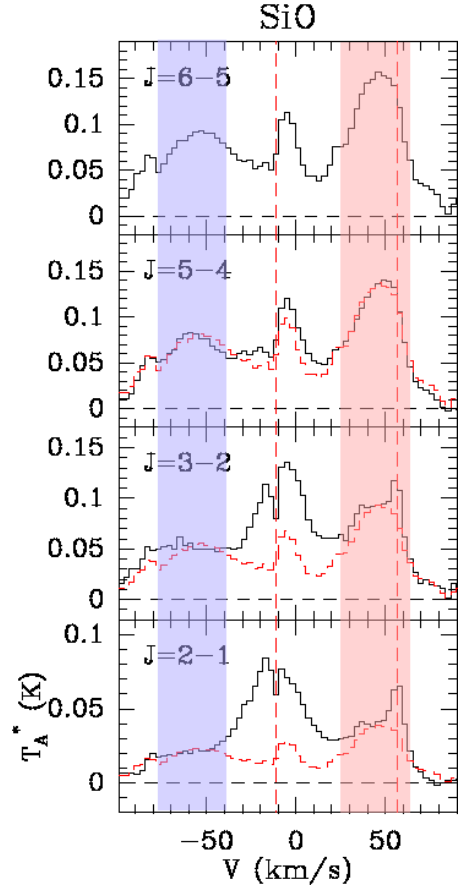


Figure 4. Montage of the SiO transitions $J=2-1$, $3-2$, $5-4$ and $6-5$, detected with the IRAM 30m telescope. The ambient cloud velocity ($v_{\text{lsr}} = -10.9 \text{ km s}^{-1}$) and the jet emission peak velocity in the northern lobe ($+56 \text{ km s}^{-1}$) are marked with red dashed lines. The SiO $J=6-5$ line profile is superimposed in dashed red on the different SiO spectra. A scaling factor has been applied to the $J=6-5$ line profile so to match the emission of the high-velocity gas in the intervals marked by red and blue filled background rectangles. Fluxes are expressed in units of antenna temperature T_A^* , corrected for atmospheric absorption.

3.2 SiO

Outflow emission was detected in all the SiO transitions present in the survey. A montage of the different transitions is displayed in Fig. A1. The jet signature in the northern lobe is unambiguously detected at $v = +56 \text{ km s}^{-1}$ in the SiO $J=2-1$ and $3-2$ transitions (Figs. 3 and 4). We failed to detect jet emission in the higher- J ($J > 3$) transitions. This is probably because the telescope beam size (HPBW) is less than $11''$ at the frequency of the high-excitation lines of SiO, hence, it encompasses a very small emitting region from the jet (see Fig. 1). Also, as can be seen in Fig. 4, the profiles of the transitions $J=5-4$ and $6-5$ show an excellent agreement once they are scaled to match the intensity of the high-velocity range. More generally, for all SiO transitions, a very good match is observed for the both blue- and redshifted gas in the velocity intervals $[-80; -40] \text{ km s}^{-1}$ and $[+25; +62] \text{ km s}^{-1}$, respectively (blue and red filled areas in Fig. 4). This suggests that the different transitions are tracing a component of high-excitation, which spans a broad

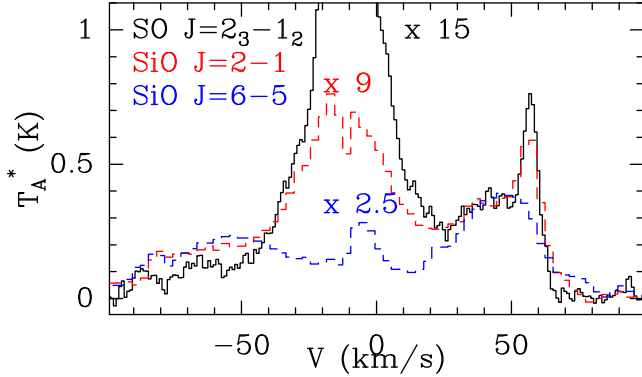


Figure 5. Comparison of the profiles of SO $J=2_3-1_2$ transition (solid black) with SiO $J=2-1$ (dashed red) and SiO $J=6-5$ (dashed blue). Magnifying factors of 15, 9 and 2.5, respectively, have been applied in order to enhance and match the high-velocity emission in the jet. Fluxes are expressed in units of antenna temperature T_A^* , corrected for atmospheric absorption.

velocity range. Since the line profile is independent of the transition and the telescope beam size, we conclude that the size of this component must be small in front of the telescope beam.

As illustrated in Fig. 4, the $J=2-1$ and $3-2$ line profiles can be decomposed as the sum of the high-excitation, broad line component traced by the $J=6-5$ and $5-4$ lines, and a "narrow" gaussian line ($\text{FWHM} \simeq 9 \text{ km s}^{-1}$), which peaks at the northern jet velocity ($+56 \text{ km s}^{-1}$). We identify the signature of the jet with this narrow line component.

3.3 SO

For SO, we detected all the transitions with an Einstein coefficient A_{ij} larger than 10^{-5} s^{-1} in the 3mm band and 10^{-4} s^{-1} in the 2mm and 1.3mm bands, respectively. We observe an excellent match in the profiles of low-excitation SO lines like $J=2_3-1_2$ or $J=3_4-2_3$ and the SiO $J=2-1$ transition (see Fig. 5 and Fig. A3), showing that SO is tracing the same physical components as SiO. In the spectra of the 3mm and 2mm transitions (left column in Fig. A3), the jet signature appears superimposed upon the emission of a broad line component. Proceeding like for the SiO analysis, we find that the jet emission can be fitted by a gaussian peaking at about $+56 \text{ km s}^{-1}$ and linewidth (FWHM) of $\approx 9 \text{ km s}^{-1}$ (see Table 1). These parameters are similar to those of the SiO jet emission.

As for the transitions in the 1.3mm band (right column in Fig. A3), a large diversity of profiles is observed, which makes it difficult to disentangle with confidence the contribution of the jet from the broad line component. The situation is more complicated than for SiO. Again, it is quite likely that the beam filling factor of the SO emission is very small and is very sensitive to calibration and pointing errors. We do not take into account these transitions. High-angular resolution with NOEMA are required to clarify this issue.

3.4 Other species

Five transitions of H_2CO (1 para- and 4 ortho-) were detected at 2mm and 1.3mm. As can be seen in Fig. A2, the

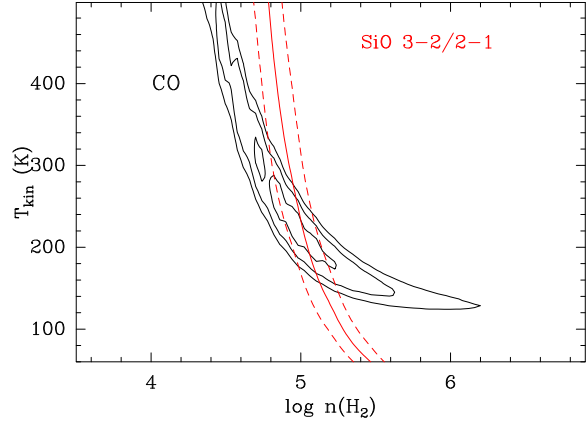


Figure 6. χ^2 distribution of LVG slab models for the CO emission in the plane $(n(\text{H}_2), T)$. Black contour levels of 0.25, 0.50, 0.75 are drawn. The solid red curve draws the observed SiO $3-2/2-1$ line intensity ratio as a function of $n(\text{H}_2)$ and T . An interval of 20% around the ratio is delineated by the dashed curves.

line profiles look all very similar. This is probably because they arise from transitions with similar Einstein coefficient A_{ij} and upper energy level E_{up} . Interestingly, the line profiles also look very different from those of SiO and SO, as can be seen in Fig. 1 for instance. The broad line component detected in the former tracers emits very weakly in the H_2CO lines, permitting disentanglement of the northern (redshifted) jet emission.

The CS $J=2-1$ and $J=3-2$ lines were detected in the jet (Fig. A2). Like for H_2CO , the northern jet emission is easily disentangled from the broad line component.

The northern jet signature could be identified in the ground state transition of HCN and HCO^+ as a weak component superimposed upon the broad line emission. Unlike the other tracers, the HCN intensity distribution decreases continuously with increasing velocity (Fig. A4).

4 DISCUSSION

4.1 Physical conditions and molecular abundances

We have constrained the physical conditions and the abundance of the detected molecular species from the modelling of the line fluxes using the radiative transfer code MADEX in the large velocity gradient (LVG) approximation (Cernicharo 2012). The adopted collisional coefficients are summarised in Table 4. For each molecular species, we have built a grid of models with density between 10^4 cm^{-3} and 10^8 cm^{-3} and temperature between 30 K and 800 K to determine the region of minimum χ^2 as a function of density and temperature.

4.1.1 The jet

In the northern jet, the parameters of the molecular transitions (flux, linewidth) were obtained from a simple Gaussian fit to the line profiles, as explained in the previous Section. The results of those fits are summarized in Table 1. In order to derive the intrinsic brightness temperatures of the molecular transitions detected in the jet,

Table 1. Observational and spectroscopic properties of the molecular transitions detected in the northern jet. The observational properties are obtained from a Gaussian fitting to the line profiles. Intensities are expressed in units of antenna temperature corrected for atmospheric absorption (T_A^*).

		ν (MHz)	E_{up} (K)	HPBW ($''$)	η_{MB}	W (K km s $^{-1}$)	V_0 (km s $^{-1}$)	Δv (km s $^{-1}$)	T_A^* (mK)	ff
(+0 $''$, +12 $''$)										
CO	1–0	115271.202	5.5	21.7	0.83	2.87(.26)	57.7(.4)	9.4(1.0)	290	0.11
	2–1 ¹	230538.000	16.5	21.7	0.64	10.49(.10)	56.0(.1)	11.5(.1)	890	0.11
	3–2 ²	345795.990	33.2	14.3	0.67	37.07(.78)	56.0(.1)	10.4(.3)	3.35×10^3	0.17
	4–3 ²	461040.768	55.3	10.7	0.45	42.77(1.70)	56.5(.2)	9.4(.4)	4.4×10^3	0.21
(+0 $''$, +0 $''$)										
CO	1–0	115271.202	5.5	21.7	0.83	2.87(.26)	57.7(.4)	9.4(1.0)	205	0.043
SiO	2–1	86846.986	6.3	28.6	0.85	0.33(.01)	56.5(.2)	9.4(.4)	35(2.7)	0.043
	3–2	130268.687	12.5	18.9	0.81	0.32(.01)	55.9(.2)	9.4(.4)	33(4.2)	0.043
	5–4	217104.920	31.3	11.3	0.65	-	-	-	< 3.9 (*)	-
SO	2 ₂ – 1 ₁	86093.957	19.3	28.6	0.85	0.102(.010)	55.5(.5)	10.9(1.4)	8.7(1.6)	0.043
	2 ₃ – 1 ₂	99299.886	9.2	24.8	0.83	0.450(.008)	56.6(.1)	8.6(.2)	49.2(1.6)	0.043
	3 ₂ – 2 ₁	109252.181	21.1	22.5	0.82	0.096(.012)	56.4(.6)	9.1(1.5)	9.9(2.0)	0.043
	3 ₃ – 2 ₂	129138.901	25.5	19.1	0.81	0.246(.015)	57.1(.2)	8.2(.6)	28.2(3.3)	0.043
	3 ₄ – 2 ₃	138178.654	15.9	17.8	0.80	0.772(.010)	56.5(.1)	8.2(.2)	88.2(2.0)	0.043
	4 ₃ – 3 ₂	158971.811	28.7	15.5	0.76	0.364(.010)	56.7(.1)	7.9(.3)	43.1(2.3)	0.043
	4 ₄ – 3 ₃	172181.403	33.8	14.3	0.74	0.346(.020)	53.6(.5)	8.8(.7)	37.1(4.0)	0.042
CS	2–1	97.980953	7.1	25.1	0.84	0.193(.009)	58.0(.2)	9.6(.5)	18.6(1.6)	0.043
	3–2	146.969026	14.1	16.7	0.78	0.223(.008)	57.1(.2)	9.4(.4)	22.3(2.6)	0.043
	5–4	244935.555	35.4	10.0	0.62	0.076(.027)	55.1(2)	5.8(3)	12.3(2.1)	0.031
H ₂ CO	2 _{1,2} –1 _{1,1}	140839.516	6.8	17.5	0.80	0.371(.008)	57.2(.1)	7.7(.2)	45.1(2.6)	0.043
	2 _{0,2} –1 _{0,1}	145602.951	10.5	16.9	0.78	0.218(.006)	57.5(.1)	7.9(.3)	25.9(2.1)	0.043
	2 _{1,1} –1 _{1,0}	150498.335	7.5	16.3	0.76	0.271(.022)	57.2(.3)	7.3(.7)	34.8(1.8)	0.043
	3 _{1,3} –2 _{1,2}	211211.449	16.9	11.6	0.66	0.202(.030)	56.3(.8)	10.9(1.9)	17.4(3.1)	0.039
	3 _{1,2} –2 _{1,1}	225697.772	18.3	10.9	0.65	0.155(.027)	55.6(.4)	8.4(1.6)	17.3(3.8)	0.036
HCO ⁺	1–0	89.188526	4.3	27.6	0.85	0.197(.015)	56.2(.3)	7.9(.7)	23.5(1.5)	0.043
HCN	1–0	88.631602	4.3	27.7	0.85	0.090(.013)	57.9(.5)	7.1(1.0)	12.1(1.3)	0.043

¹Convolved at the resolution of the CO $J=1-0$ line.²Observations with the JCMT.

the main-beam brightness temperature obtained from our simple Gaussian fit analysis was corrected for the coupling between the jet and the telescope main-beam. The latter was estimated following the approach described in Lefloch et al. (2015) in their analysis of the southern lobe of the CO jet. In a first step, the solid angle of the jet is estimated from the intersection between the half-power contour of the CO distribution mapped with the PdBI and the telescope main-beam. The latter is modeled by a disk centered at the position of the protostar ($0'', 0''$), of diameter equal to the telescope HPBW. In a second step, the corresponding beam filling factor is obtained from the ratio of the source solid angle to the telescope main beam. In our LVG calculations, we have adopted a linewidth $\Delta v = 10 \text{ km s}^{-1}$ for CO, CS and SiO, and $\Delta v = 8 \text{ km s}^{-1}$ for H₂CO and SO, in agreement with the results of our Gaussian fits to the line profiles (Table 1). Our PdBI CO map shows that the northern jet emission is detected from a few arcsec away from the proto-

star (Fig. 1). The pointing uncertainties for the IRAM 30m telescope, which are typically $2''$ – $3''$, cast some uncertainties on the actual beam filling factor for the observations in the 1.3mm band, when the telescope beam (HPBW) is $\approx 10''$ and the covered jet length becomes smaller than the pointing uncertainties. For this reason, we have considered only the millimeter transitions detected in the 3mm and 2mm bands.

In the southern jet, molecular line profiles, like e.g. SiO $J=2-1$, deviate strongly from gaussianity (see also Fig. 3). For this reason, we have taken into account only the flux integrated in the velocity range between -150 and -100 km s^{-1} . Like for the northern jet, we have considered only the molecular transitions detected in the 3mm and 2mm bands, and for each species we have determined the molecular gas column density adopting the kinetic temperature derived from the CO multi-transition analysis of Lefloch et

Table 2. Velocity-integrated flux and spectroscopic properties of the molecular transitions detected at 3mm and 2mm in the southern jet. Line fluxes W are integrated between -150 and -100 km s $^{-1}$. Intensities are expressed in units of antenna temperature corrected for atmospheric absorption (T_A^*).

		ν (MHz)	E_{up} (K)	HPBW ($''$)	η_{MB}	W (K km s $^{-1}$)	ff
CO	1-0	115271.202	5.5	21.7	0.83	0.67	0.043
SiO	2-1	86846.986	6.3	28.6	0.85	0.10	0.043
	3-2	130268.687	12.5	18.9	0.81	0.18	0.043
SO	2 ₂ - 1 ₁	86093.957	19.3	28.6	0.85	0.03	0.043
	2 ₃ - 1 ₂	99299.886	9.2	24.8	0.83	0.04	0.043
	3 ₂ - 2 ₁	109252.181	21.1	22.5	0.82	-	0.043
	3 ₃ - 2 ₂	129138.901	25.5	19.1	0.81	-	0.043
	3 ₄ - 2 ₃	138178.654	15.9	17.8	0.80	0.05	0.043
	4 ₃ - 3 ₂	158971.811	28.7	15.5	0.76	0.04	0.043
	4 ₄ - 3 ₃	172181.403	33.8	14.3	0.74	0.04	0.042
CS	2-1	97.980953	7.1	25.1	0.84	-	0.043
	3-2	146.969026	14.1	16.7	0.78	-	0.043
H ₂ CO	2 _{1,2} -1 _{1,1}	140839.516	6.8	17.5	0.80	-	0.043
	2 _{0,2} -1 _{0,1}	145602.951	10.5	16.9	0.78	-	0.043
	2 _{1,1} -1 _{1,0}	150498.335	7.5	16.3	0.76	-	0.043
HCO ⁺	1-0	89.188526	4.3	27.6	0.85	0.39	0.043
HCN	1-0	88.631602	4.3	27.7	0.85	0.02	0.043

al. (2015). The velocity-integrated fluxes are reported in Table 2.

Our best fit model of the CO flux in the transitions $J=1-0$ to $4-3$ yields a temperature $T_{\text{kin}} \approx 220$ K, a gas column density $N(\text{CO}) = 2.7 \times 10^{17}$ cm $^{-2}$, and a gas density $n(\text{H}_2) \approx 1.0 \times 10^5$ cm $^{-3}$. We report in Fig. 6 our χ^2 analysis of the solutions in the plane (n, T) . Degeneracy is observed and possible solutions are found in the range $T_{\text{kin}} = 180-300$ K and $n(\text{H}_2) = (0.6-2.0) \times 10^5$ cm $^{-3}$. Those observations were carried out towards the offset position ($0''$, $+12''$) but the beam size is large enough to encompass material close to the protostar (Figs. 1-2). These physical conditions should represent a good approximation to those at the origin of the jet. Indeed, our radiative transfer code predicts a CO $J=1-0$ line intensity in reasonable agreement with the observations, when taking into account the jet-beam coupling. In what follows, we will adopt these values for the jet physical parameters in the direction of the protostar. They are summarized in Table 5.

Analysis of the SO emission brings a very good constraint on the H₂ gas density, with $n(\text{H}_2) = (1-2) \times 10^5$ cm $^{-3}$. Analysis of the CS, SiO and H₂CO lines are consistent with these results. This is illustrated in Fig. 6, in which we have reported the value of the SiO 3-2/2-1 line intensity ratio in the parameter space (n, T) . This ratio is a good probe of $n(\text{H}_2)$, whereas it is barely sensitive to the kinetic temperature. In the case of HCN and HCO⁺, for which only one transition was detected, we ran our LVG code to determine the gas column density, adopting a kinetic temperature $T=220$ K and a gas density $n(\text{H}_2)=10^5$ cm $^{-3}$.

The derived column densities are summarized in Table 6. The molecular abundances with respect to H₂ were obtained from the ratio of the derived column densities to that of CO, assuming a standard abundance ratio $[\text{CO}]/[\text{H}_2]=10^{-4}$. They are summarized in Table 7.

In the southern jet, the gas density $n(\text{H}_2)$ derived from the SiO lines is of the order of $(2-4) \times 10^5$ cm $^{-3}$. An upper limit was estimated for the species missed in the southern jet: HCN, CS and H₂CO. The results are summarized in Tables 5-7. Despite the difference in column densities between the northern and southern jet, it appears that all the molecular species display similar abundances, of the order of 10^{-8} with respect to H₂. This point is further addressed in the next section.

4.1.2 The low-velocity outflow

We have estimated the molecular abundances of the various jet tracers in both low-velocity outflow lobes following the same methodology as for the jet. In order to avoid contamination of the outflow emission by the hot corino whose line profiles have a typical linewidth of $5-6$ km s $^{-1}$ (see Ospina-Zamudio et al. 2018) and peak at the ambient cloud velocity $v_{\text{lsr}} = -10.9$ km s $^{-1}$, the outflow molecular line fluxes were measured in the velocity intervals $[-20;-14]$ km s $^{-1}$ and $[-8;-2]$ km s $^{-1}$ for the southern and northern lobes, respectively. The velocity-integrated fluxes measured in both outflow lobes are summarized in Table 3.

The filling factor was obtained from computing the convolution of the telescope main beam with the brightness dis-

Table 3. Velocity-integrated emission of the jet tracers in the northern and southern low-velocity outflow cavity lobes. The fluxes are integrated in the velocity intervals $[-20; -14]$ km s^{-1} and $[-8; -2]$ km s^{-1} , associated with the southern and northern outflow lobes, respectively. Intensities are expressed in units of antenna temperature corrected for atmospheric absorption (T_A^*).

		ν (MHz)	E_{up} (K)	HPBW ($''$)	η_{MB}	W($[-20; -14]$) (K km s^{-1})	W($[-8; -2]$) (K km s^{-1})	ff
CO	1-0	115271.202	5.5	21.7	0.83	11.9	9.8	0.20
	2-1	230538.000	16.5	21.7	0.64	23.2	20.7	0.20
	3-2	345795.990	33.2	14.3	0.67	27.2	25.0	0.24
SiO	2-1	86846.986	6.3	28.6	0.85	0.48	0.44	0.17
	3-2	130268.687	12.5	18.9	0.81	0.52	0.75	0.21
	5-4	217104.920	31.3	11.3	0.65	0.38	0.71	0.26
	6-5	260518.018	43.8	9.4	0.58	0.31	0.62	0.28
	7-6	303926.814	58.4	8.1	0.52	0.21	0.87	0.29
	8-7	347330.592	75.0	7.1	0.43	0.18	0.41	0.31
SO	$2_2 - 1_1$	86093.957	19.3	28.6	0.85	0.11	0.15	0.17
	$2_3 - 1_2$	99299.886	9.2	24.8	0.83	0.78	0.69	0.18
	$3_2 - 2_1$	109252.181	21.1	22.5	0.82	0.19	0.15	0.19
	$3_3 - 2_2$	129138.901	25.5	19.1	0.81	0.26	0.34	0.21
	$3_4 - 2_3$	138178.654	15.9	17.8	0.80	0.89	1.20	0.22
	$4_3 - 3_2$	158971.811	28.7	15.5	0.76	0.32	0.52	0.23
	$4_4 - 3_3$	172181.403	33.8	14.3	0.74	0.31	0.56	0.24
	$5_4 - 4_3$	206.176005	38.6	11.9	0.67	0.45	0.67	0.26
	$5_5 - 4_4$	215.220650	44.1	11.4	0.66	0.27	0.86	0.26
	$5_6 - 4_5$	219.949388	35.0	11.2	0.66	0.78	1.25	0.27
	$6_5 - 5_4$	251.825759	50.7	9.8	0.60	0.24	0.89	0.28
	$6_6 - 5_5$	258.255826	56.5	9.5	0.58	0.20	0.72	0.29
	$6_7 - 5_6$	261.843705	47.6	9.4	0.58	0.71	1.16	0.29
CS	2-1	97.980953	7.1	25.1	0.84	1.50	0.96	0.14
	3-2	146.969026	14.1	16.7	0.78	1.59	1.58	0.18
	5-4	244935.555	35.4	10.0	0.62	0.85	1.51	0.22
	6-5	293912.089	49.4	8.4	0.52	0.53	2.22	0.24
	7-6	342882.854	65.8	7.2	0.43	0.41	0.96	0.25
H ₂ CO	$2_{1,2}-1_{1,1}$	140839.516	6.8	17.5	0.80	1.89	1.88	0.17
	$2_{0,2}-1_{0,1}$	145602.951	10.5	16.9	0.78	1.47	1.13	0.18
	$2_{1,1}-1_{1,0}$	150498.335	7.5	16.3	0.76	1.77	1.51	0.18
	$3_{1,3}-2_{1,2}$	211211.449	16.9	11.6	0.66	1.93	1.74	0.21
	$3_{1,2}-2_{1,1}$	225697.772	18.3	10.9	0.65	1.33	1.80	0.22
HCO ⁺	1-0	89.188526	4.3	27.6	0.85	1.88	1.13	0.17
HCN	1-0	88.631602	4.3	27.7	0.85	4.37	3.65	0.17

Table 4. Collisional coefficients adopted for LVG calculations with MADEX

Species	Collisional coefficients
CO	CO-H ₂ Yang et al. (2010)
SiO	SiO-p-H ₂ Dayou & Balança (2006)
SO	SO-He Lique et al. (2006)
H ₂ CO	H ₂ CO-H ₂ from Green (1991)
CS	CS-He Lique et al. (2007)
HCN	HCN-He Dumouchel et al. (2010)
HCO ⁺	HCO ⁺ -H ₂ Flower et al. (1999)

tribution of the outflow cavity. We assumed that the latter is uniform and the shape of the outflow cavity close to the

protostar can be modelled as a parabola $z = 0.44r^2$ where the height z and radius r are expressed in arcsec units (see Fig. B1).

Our LVG analysis of the CO transitions $J=1-0$, $2-1$ and $3-2$ indicates similar physical conditions in both lobes with a gas kinetic temperature $T_{\text{kin}} \simeq 50$ K and a column density $N(\text{CO}) \simeq 1.2 \times 10^{17} \text{ cm}^{-2}$. Multi-line analysis of SO and CS yield consistent, typical $n(\text{H}_2)$ densities in the range $(1-4) \times 10^5 \text{ cm}^{-3}$ and $(1-8) \times 10^5 \text{ cm}^{-3}$ in the northern and southern outflow lobes, respectively. These results are consistent with the physical conditions previously estimated towards the southern outflow lobe by Lefloch et al. (2015).

In the case of the SiO emission, it is not possible to fit the whole spectral line energy distribution (SLED) with one single gas component. We note that Lefloch et al. (2015)

Table 5. Physical conditions determined in the northern and southern lobes for the high-velocity jet and the low-velocity outflow from a LVG analysis of the CO emission with the MADEX radiative transfer code.

	Jet		Outflow	
	North	South	North	South
Vel. range (km s ⁻¹)	+50;+80	-150;-100	-8;-2	-20;-14
T_{kin} (K)	180-300	80-100	50	50
$n(\text{H}_2)$ (10 ⁵ cm ⁻³)	0.6-2.0	0.5-1.0	1-4	1-7
$N(\text{CO})$ (10 ¹⁷ cm ⁻²)	2.7	0.9	1.2	1.4

Table 6. Molecular gas column densities determined in the northern and southern lobes for the high-velocity jet and the low-velocity outflow from a LVG analysis with MADEX.

Species	Jet		Outflow	
	North	South	North	South
	N (cm ⁻²)	N (cm ⁻²)	N (cm ⁻²)	N (cm ⁻²)
CO	2.7(17)	0.9(17)	1.2(17)	1.4(17)
SiO	1.7(13)	1(13)	6(12)	5.5(12)
SO	1.0(14)	1(13)	8.5(13)	5.5(13)
H ₂ CO	1.2(14)	< 2(12)	7.0(13)	8.1(13)
CS	3.3(13)	< 3(12)	3.7(13)	5.5(13)
HCN	3.3(12)	< 5(11)	(4.0-7.5)(13)	1.0(14)
HCO ⁺	4.7(12)	1(13)	(5.5-8.0)(12)	1.7(13)

reported evidence for a gradient of excitation conditions in the outflow (see their Table 4). Therefore, we have modelled only the emission of the lower-excitation SiO lines $J=2-1$ and $3-2$. We obtained typical $n(\text{H}_2)$ densities of a few 10⁵ cm⁻³, consistent with the SO and CS analysis.

In the case of H₂CO, we modelled separately the emission of the ortho- para- lines. All four ortho transitions $2_{1,2}-1_{1,1}$, $2_{1,1}-1_{1,0}$, $3_{1,3}-2_{1,2}$ and $3_{1,2}-2_{1,1}$ were fitted simultaneously. We then adopted the same kinetic temperature and H₂ density and we determined the column density $N(\text{p-H}_2\text{CO})$ necessary to reproduce the flux of the para transition $2_{0,2}-1_{0,1}$. We obtain an ortho to para ratio of 2.5, which is compatible with the canonical value of 3, taking into account the flux uncertainties.

The outflow physical conditions, molecular gas column

Table 7. Molecular abundances with respect to H₂ determined in the northern and southern lobes for the high-velocity jet and the low-velocity outflow, assuming a standard abundance ratio $[\text{CO}]/[\text{H}_2] = 10^{-4}$.

Species	Jet		Outflow	
	North	South	North	South
	X	X	X	X
SiO	0.6(-8)	1(-8)	0.5(-8)	0.4(-8)
SO	3.7(-8)	1(-8)	7.1(-8)	3.9(-8)
H ₂ CO	4.4(-8)	< 0.2(-8)	5.8(-8)	5.8(-8)
CS	1.2(-8)	< 0.3(-8)	3.1(-8)	3.9(-8)
HCN	0.12(-8)	< 0.05(-8)	(3.3-6.3)(-8)	7.1(-8)
HCO ⁺	0.17(-8)	1(-8)	(0.46-0.67)(-8)	1.2(-8)

densities and abundances are summarized in Tables 5, 6 and 7, respectively.

4.2 Chemical differentiation

How different are the physical and chemical conditions between the northern and the southern jets? The physical conditions estimated in the northern jet indicate a higher temperature ($T \simeq 200$ K) and a higher CO column density ($N(\text{CO}) \simeq 3 \times 10^{17}$ cm⁻²) than in the southern lobe ($T \simeq 100$ K and $N(\text{CO}) \simeq 1 \times 10^{17}$ cm⁻²). The high-velocity CO gas distribution, as observed with the PdBI (Fig. 1) displays the same "knotty" structure in both lobes. Lefloch et al. (2015) interpreted these knots as the signature of internal shocks which act as a local source of gas heating in the jet, and are responsible for the high temperature (750–1200 K), dense ($\sim 10^6$ cm⁻³) CO gas component detected in the southern lobe. We speculate that similar excitation components are also present in the northern jet.

The jet dimensions (length and transverse size) as determined from the brightness contour at half power are found to be similar in both the northern and southern jets, and can be fitted approximately by an ellipse of $20'' \times 1.7''$, which implies a typical radius of $0.85''$ or 620 au for the jet. The radius of the high-velocity CO jet does not display significant variations with distance to the protostar and is ~ 6 times larger than the values reported towards molecular jets of low-mass protostars like HH212 (90AU, Cabrit et al. 2012) or L1157 (60-125AU, Podio et al. 2016). The CO brightness distribution appears rather loose at low intensity levels, suggesting the presence of a complex structure at subarcsec scale. The Parallaxic Angle is different between the two lobes however, with values of $+64^\circ$ and $+69^\circ$, respectively. We note that the gas distribution extends to larger radii, up to about $1.7''$ ($2.0''$) in the southern (northern) jet, which is suggestive of *entrained gas* along the jet.

In the jet, a comparison of the molecular abundances in the northern and southern lobe shows that they are similar within a factor of 2. Differences in the relative composition are detected, however: the most abundant species in the northern jet are SO and H₂CO, with an abundance of about 4×10^{-8} relative to H₂; on the contrary, the most abundant species in the southern jet are SO and SiO, with a relative abundance of 2×10^{-8} . Also, HCO⁺, appears to be more abundant by a factor of 2 in the southern jet. CS and H₂CO, which are both undetected in the southern jet, are underabundant by a factor of 4 *at least* with respect to the northern jet.

In the outflow, our analysis shows that the physical and chemical conditions in the northern and southern lobes are also very similar. As a matter of fact, the abundances of the different tracers considered in the present work (i.e. those detected in the jet) vary by less than a factor of 2 between the two outflow lobes (see Table 7). Our unbiased spectral line survey of CepE-mm shows that the jet is often not detected in the tracers of the chemically-rich outflow gas, which proves a strong chemical differentiation between the jet and the outflow. As an example, none of the complex organic molecules (COMs) detected in the low-velocity outflow (the *extremely broad line component* eBL in Ospina-Zamudio et al. 2018) displays any detectable signature in the protostellar jet.

It is rather surprising that molecular species such as H_2CO and CS , which are detected in both outflow lobes with similar abundances, are detected only in the northern jet, with an abundance similar to that of the outflow gas. Conversely, the abundances of molecular species detected both in the outflow and the jet are similar within a factor of a few (Table 7). The only exception is HCN , which is found much lower in the jet by one order of magnitude. We must take care in not overinterpreting these quantitative results as we cannot exclude that the emission of the different tracers may arise from different regions of the jet. The question arises, however, as to whether such a chemical differentiation is caused by the different physical conditions between the jet and the outflow or by a fundamentally different origin between the high-velocity material and the low-velocity outflow gas.

4.3 Origin of the emission

4.3.1 Evidence for jet entrained gas

Analysis of the molecular line emission reveals the presence of several species with abundances of the order of 10^{-8} . The detection of these various molecular species makes Cep E-mm one of the richest molecular jets reported until now, together with L1448 and I04166 (Tafalla et al. 2010). Several observational facts are intriguing and are difficult to reconcile with an origin of the molecular emission in the protostellar jet. First, the jet diameter derived from the CO emission is about one order of magnitude larger than the values measured in typical low-mass Class 0 jets or predicted by models, which are usually of the order of a few times 10 au. This suggests that CO is not tracing the protostellar jet alone. Second, some of the most abundant species, like H_2CO (4.4×10^{-8} in the northern jet) or HCO^+ (10^{-8} in the southern jet) are usually not detected in young protostellar jets. Last, a strong degree of chemical differentiation is observed between the northern and southern jets, which is difficult to understand if the jet is made of pristine material taken away from the protostellar disk or formed in the jet itself.

Therefore, it is quite likely that the emission detected in our spectral survey does *not* originate from the protostellar jet itself. We note that both the jet and the low-velocity outflow are dense, with a density $n(\text{H}_2) \approx$ a few times 10^5 cm^{-3} in the jet up to $\sim 10^6 \text{ cm}^{-3}$ in the entrained gas of the outflow. Under such conditions, the formation of a turbulent molecular layer at the interface between the jet and outflow cavity gas is to be expected.

The results of our study are biased by the lack of angular resolution of the IRAM 30m spectral line survey and we cannot exclude that different species are actually tracing different regions of the jet/outflow system. This hypothesis is supported by the differences observed between the line profiles of the molecular tracers (see Fig. 1). Interferometric observations at arcsec scale are required to clarify this point and better understand the chemical differentiation along the jet.

4.3.2 Shock-driven photochemistry at work?

Several of the measured molecular abundances are similar in the jet and the outflow and it is not clear whether this similarity is coincidental or it reflects the nature of the molecular-rich high-velocity gas, i.e. entrained gas from the outflow cavity. Even so, this hypothesis does not explain why some species are present in one lobe and not the other, as illustrated by HCO^+ and H_2CO . Conversely, the HCN abundance in the jet is about one order of magnitude less than in the low-velocity outflow.

We propose that the unusual chemistry observed in/around the Cep E-mm could be driven by internal shocks along the jet or in the turbulent layer of outflow gas entrained by the protostellar jet. As discussed by Lefloch et al. (2015) and in this work, there is strong evidence for the propagation of shocks and instabilities along the jet. The dynamical age of the jet is short enough, of the order of 1000 yr, that shocks are expected to display mixed C- and J-type components. We speculate that these shock components create a diffuse UV radiation field which will also illuminate the surrounding outflow gas.

Viti et al. (2002) have investigated the impact of such a shock-driven UV field on the chemical composition of the parental cloud gas at the interface with the outflow. Their modelling cannot be directly applied to the present situation as the chemical and physical conditions in the outflow gas are different from the quiescent envelope considered by them. On the one hand, the outflow gas is warmer than the envelope, and, on the other hand, dust grains have already been processed and a fraction of their icy mantles eroded through the bowshock responsible for the formation of the outflow. However, this modelling can bring some insight on processes at work in the Cep E-mm outflow/jet system. In their modelling, Viti et al (2002) consider a shock producing a UV radiation field with an intensity of $\chi = 10$ times the Habing radiation field (Habing, 1968). The shock-induced photoionization drives a rich photochemistry in the immediate vicinity of the outflow and progresses into the core material. As can be seen in their Figure 2, H_2CO and SO are produced with similarly high abundances ($\sim 10^{-7}$) and remain for several 10 yr in the gas phase. Interestingly, they find that right after CO and H_2O , the most abundant species is HCO^+ (abundances \sim a few 10^{-9}), provided that the surrounding gas temperature is not too high (30 – 100 K). A higher temperature (~ 1000 K) extends the duration of the chemically-rich phase, resulting in HCO^+ abundance of $\sim 10^{-8}$. Interestingly, the predictions of this model appear to agree qualitatively with our observations. This agreement is very encouraging and supports the scenario of a shock-driven chemistry in a turbulent layer around the high-velocity jet. Interferometric observations at subarcsec scale are required to elucidate the chemical differentiation observed along the jet and the behaviour of the different tracers, so to confirm or rule out the role of shock-driven chemistry in the outflow.

5 CONCLUSIONS

We have carried out a systematic search for molecular line emission from the Cep E-mm high-velocity protostellar jet using the IRAM 30m telescope, complemented

with JCMT and PdBI observations. In addition to CO, we have identified several molecular species in the jet: SO, SiO, H₂CO, CS, HCO⁺ and HCN. We find evidence for chemical differentiation between the northern and southern jet lobes. Whereas all these aforementioned species are detected in the northern jet, only CO, SO, SiO and HCO⁺ are detected in the southern jet.

A radiative transfer analysis of the molecular line emission in the Large Velocity Gradient approximation with MADEX has permitted us to constrain the physical and chemical conditions in the jet and the low-velocity outflow. We find the northern jet is made of hot (180 – 300 K), dense $((0.6-2.0) \times 10^5 \text{ cm}^{-3})$ gas. Overall, the jet material appears warmer and denser in the northern than in the southern jet, where a kinetic temperature of 80–100 K and density of $(0.5-1.0) \times 10^5 \text{ cm}^{-3}$ were determined by Lefloch et al. (2015). Also, the jet appears warmer and less dense than the entrained outflowing gas, with $T_{\text{kin}} \simeq 50 \text{ K}$ and $n(\text{H}_2) = (1-7) \times 10^5 \text{ cm}^{-3}$. Interestingly, the physical conditions in the low-velocity gas are very similar in both outflow lobes.

The molecular abundances derived in the jet are of the order of 10^{-8} , and are found to be similar to those determined in the low-velocity outflow gas. They reveal an unusual chemical composition with respect to that commonly observed in protostellar jets. A strong chemical differentiation is observed between the two jet lobes. In the northern jet, we find that H₂CO and SO are the most abundant species, immediately after CO. In the southern jet, the most abundant species are SO, SiO and HCO⁺.

The transverse size of the CO jet was estimated $\approx 1200 \text{ au}$ from Plateau de Bure observations at 1'' angular resolution. This is much larger than the values reported in protostellar jets and suggests that the jet entrains a turbulent gas envelope in the course of its propagation.

A qualitative agreement is observed between the jet chemical composition and the predictions of a simple model of a turbulent layer at the outflow/cloud interface by Viti et al. (2002). We propose that the detected molecular species are the signatures of the specific photochemistry driven by the UV field generated by shocks in the turbulent envelope.

Interferometric observations at arcsec scale are required to elucidate the origin of the jet molecular emission and to understand the observed differentiation in the protostellar gas.

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APPENDIX A: MOLECULAR TRANSITIONS DETECTED IN THE PROTOSTELLAR JET OF CEP E-MM.

Fig. A1–A4 present panels of all the detected transitions of SiO, CS, H₂CO, SO, HCN and HCO⁺ in the protostellar jet of Cep E-mm.

APPENDIX B: MODELLING OF THE OUTFLOW CAVITIES

Modelling of the outflow cavities brightness distribution was performed from the PdBI map in the CO $J=2-1$ line. We assumed that the outflow brightness distribution is uniform and that the shape of the cavities close to the protostar can be modelled as a parabola $z = 0.44r^2$ where the height z and radius r are expressed in arcsec units (see Fig. B1).

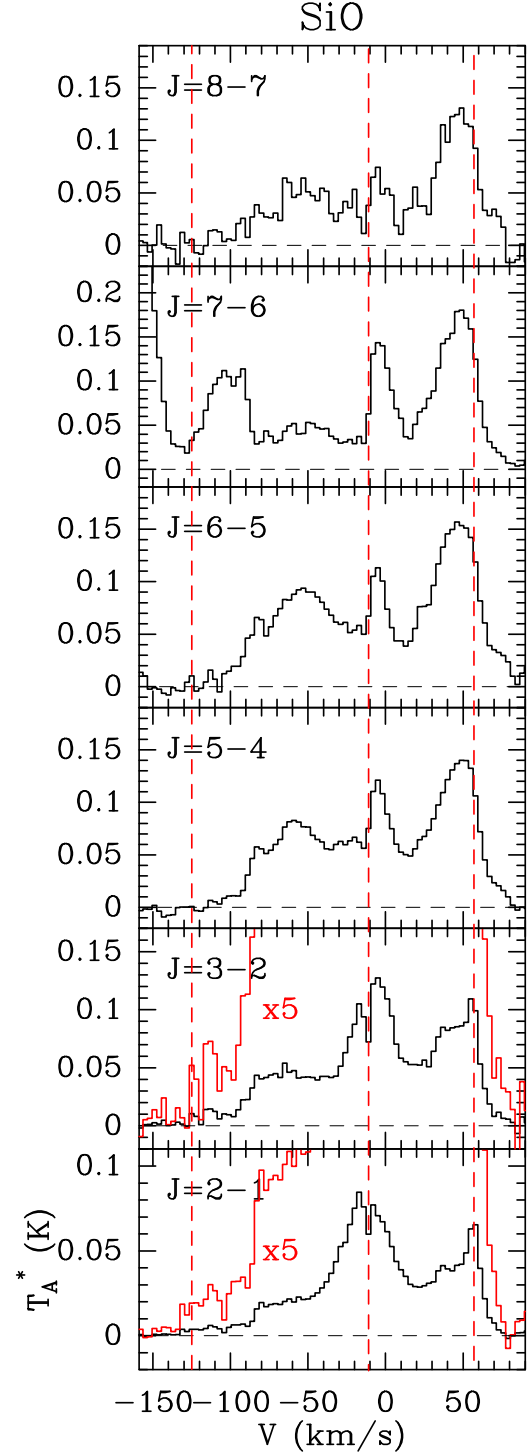


Figure A1. Montage of the SiO transitions detected with the IRAM 30m telescope. The jet emission peak velocity in the blue and redshifted lobe at $v_{lsr} = +57 \text{ km s}^{-1}$ and -125 km s^{-1} , respectively, is marked with a red dashed line. Fluxes are expressed in units of antenna temperature T_A^* , corrected for atmospheric absorption.

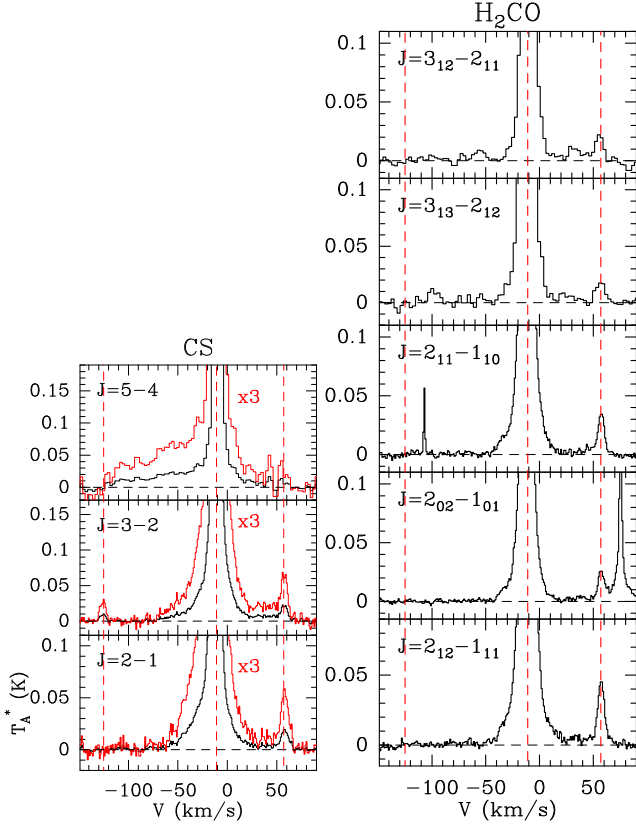


Figure A2. (left) Montage of the CS transitions (left) and H_2CO transitions (right) detected with the IRAM 30m telescope. The jet emission peak velocity in the blue- and redshifted lobe, at $v_{lsr} = -125 \text{ km s}^{-1}$ and $+57 \text{ km s}^{-1}$, respectively, is marked with a red dashed line. Fluxes are expressed in units of antenna temperature T_A^* , corrected for atmospheric absorption.

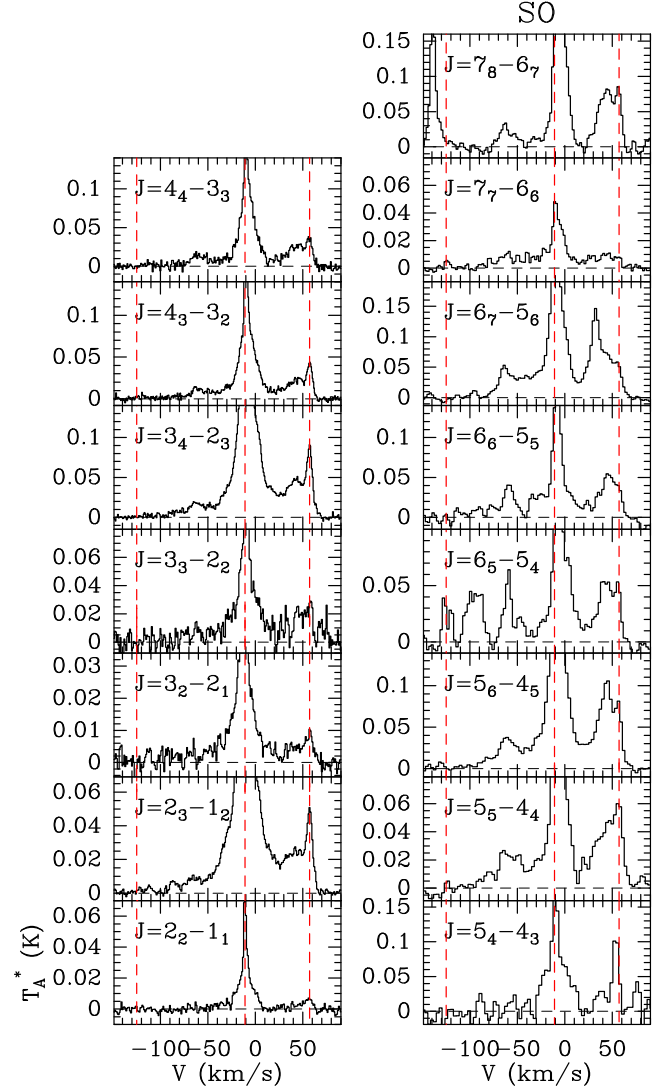


Figure A3. Montage of the SO transitions detected with the IRAM 30m telescope. The jet emission peak velocity in the blue- and redshifted lobe, at $v_{lsr} = -125 \text{ km s}^{-1}$ and $+57 \text{ km s}^{-1}$, respectively, is marked with a red dashed line. Fluxes are expressed in units of antenna temperature T_A^* , corrected for atmospheric absorption.

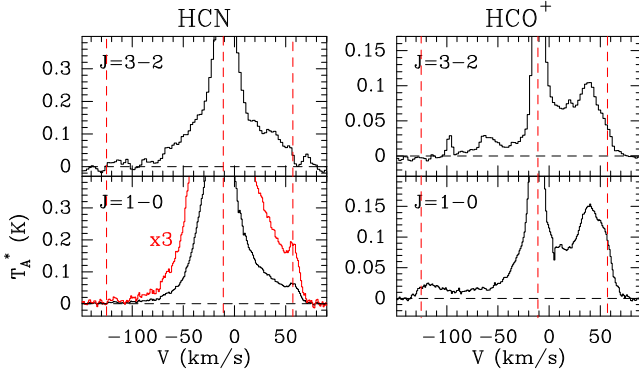


Figure A4. (*Left*) Montage of the HCN transitions detected with the IRAM 30m telescope. The jet emission peak velocity in the blue- and redshifted lobe, at $v_{lsr} = -125 \text{ km s}^{-1}$ and $+57 \text{ km s}^{-1}$, respectively, is marked with a red dashed line. (*Right*) Montage of the HCO^+ transitions detected with the IRAM 30m telescope. The jet emission peak velocity in the blue- and redshifted lobe, at $v_{lsr} = -125 \text{ km s}^{-1}$ and $+57 \text{ km s}^{-1}$, respectively, is marked with a red dashed line. Fluxes are expressed in units of antenna temperature T_A^* , corrected for atmospheric absorption.

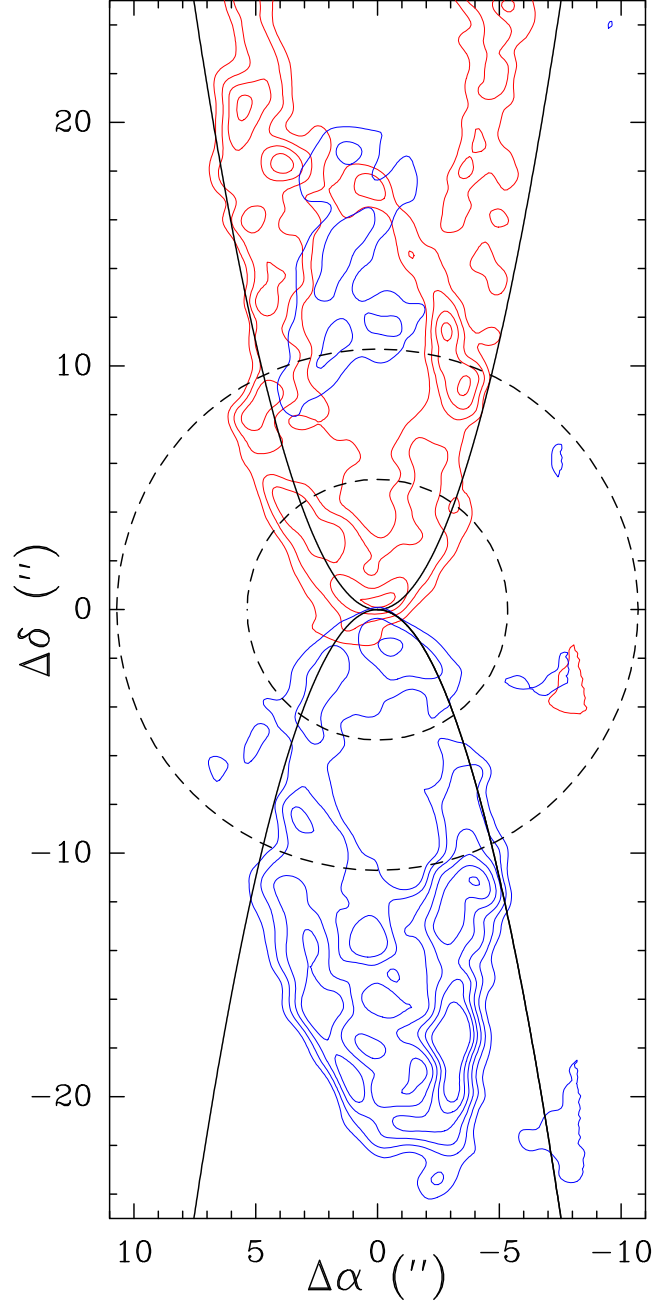


Figure B1. CO $J=2-1$ emission in the low velocity outflow cavity as observed with the Plateau de Bure interferometer at $1''$ resolution. The emission is displayed in a rotated frame so that the main axis of each outflow cavity coincides with the δ -axis. In the northern lobe (red contours) the flux is integrated between -8 and -2 km s^{-1} and the rotation angle is 12° . In the southern lobe (blue contours), the flux is integrated between -20 and -14 km s^{-1} and the rotation angle is 20° . First contour and contour interval are $1.3 \text{ Jy/beam km s}^{-1}$ in both lobes. The fit to the outflow cavities $\delta = \pm 0.44\alpha^2$ is superimposed in black. The size of the IRAM 30m main beam (HPBW) at the frequency of the CO $J=1-0$ and $2-1$ lines drawn by the circles in dashed.