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Correlation of gas dynamics and dust in the evolved filament G82.65-02.00^{★,★★}

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ABSTRACT

Context. The combination of line and continuum observations can provide vital insight into the formation and fragmentation of filaments and the initial conditions for star formation. We have carried out line observations to map the kinematics of an evolved, actively star forming filament G82.65-2.00. The filament was first identified from the *Planck* data as a region of particularly cold dust emission and was mapped at 100–500 μm as a part of the *Herschel* key program Galactic Cold Cores. The *Herschel* observations cover the central part of the filament, corresponding to a filament length of ~ 12 pc at the assumed distance of 620 pc.

Aims. CO observations show that the filament has an intriguing velocity field with several velocity components around the filament. In this paper, we study the velocity structure in detail, to quantify possible mass accretion rate onto the filament, and study the masses of the cold cores located in the filament.

Methods. We have carried out line observations of several molecules, including CO isotopologues, HCO⁺, HCN, and CS with the Osaka 1.85 m telescope and the Nobeyama 45 m telescope. The spectral line data are used to derive velocity and column density information.

Results. The observations reveal several velocity components in the field, with strongest line emission concentrated to velocity range $\sim [3, 5]$ km s⁻¹. The column density of molecular hydrogen along the filament varies from 1.0 to 2.3×10^{22} cm². We have examined six cold clumps from the central part of the filament. The clumps have masses in the range 10–20 M_{\odot} ($\sim 70 M_{\odot}$ in total) and are close to or above the virial mass. Furthermore, the main filament is heavily fragmented and most of the substructures have a mass lower than or close to the virial mass, suggesting that the filament is dispersing as a whole. Position-velocity maps of ¹²CO and ¹³CO lines indicate that at least one of the striations is kinematically connected to two of the clumps, potentially indicating mass accretion from the striation onto the main filament. We tentatively estimate the accretion rate to be $\dot{M} = 2.23 \times 10^{-6} M_{\odot}/\text{yr}$.

Conclusions. Our line observations have revealed two or possibly three velocity components connected to the filament G82.65-2.00 and putative signs of mass accretion onto the filament. The line observations combined with *Herschel* and WISE maps suggest a possible collision between two cloud components.

Key words. ISM: clouds – dust, extinction – ISM: structure – ISM: kinematics and dynamics – ISM: lines and bands

1. Introduction

Filamentary structures have been found to be common features of interstellar molecular clouds. Early results from the *Herschel* space observatory have shown the abundance of ~ 1 parsec-scale structures even in the diffuse and non-star-forming clouds

* *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

** The reduced data cubes are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/608/A21>

(André et al. 2010; Arzoumanian et al. 2011; Men'shchikov et al. 2010; Juvela et al. 2012; Rivera-Ingraham et al. 2016). Based on these observational results, it has been suggested that filamentary structures form as a result of the dissipation of large-scale turbulence and not as a result of star formation (Arzoumanian et al. 2011). On the other hand, the filamentary structures have been found to be closely connected to star formation, especially in the nearby fields where the *Herschel* resolution is sufficient to resolve gravitationally bound cores. The close relationship between cores and filaments lead Men'shchikov et al. (2010) and Könyves et al. (2010) to suggest that core formation follows a two-step path. The large-scale

turbulence and magnetic fields generate a complex web of filamentary structures followed by gravitational fragmentation of the densest structures.

Understanding the morphology and fragmentation of filamentary structures is vital in order to understand star formation (e.g. Inutsuka & Miyama 1992; Kainulainen et al. 2016). On the other hand, the internal motions of filaments and interactions between the filament and surrounding medium can provide insight into the evolution from cores to stars (Miettinen & Offner 2013). In the L1517 and B213 clouds in the Taurus cloud complex, a number of velocity-coherent structures were identified from $C^{18}O$ and N_2H^+ data by Hacar & Tafalla (2011) and Hacar et al. (2013). The coherent nature of the structures was interpreted as evidence of physically distinct sub-filaments. The authors suggested that the structures formed at different times and only some of them are producing dense cores. Furthermore, the filaments identified in the nearby cloud L1517 were found to be surrounded by subsonic gas (Hacar & Tafalla 2011), suggesting that the formation of the velocity-coherent structures occurred before fragmentation. The formation of velocity coherent substructures has been connected to the evolutionary stage of the filament. In their recent paper, Cox et al. (2016) studied the Musca filament and compared its structure with the B211/213 filament system. Both of the filaments are quite similar with a main filament 10 pc in length surrounded by a number of fainter and roughly perpendicular substructures. However, the B211/213 filament has complex internal-velocity coherent structures while there is significantly less substructure in Musca. The differences led Cox et al. (2016) to suggest that the Musca cloud represents an earlier evolutionary stage.

The faint substructures, striations seen around the filaments, can be related to mass accretion from the surrounding medium on the main filament. The findings of Arzoumanian et al. (2013) and Palmeirim et al. (2013) point to a scenario where filaments continue to accrete material throughout their evolution while retaining approximate virial balance. However, for non-equilibrium filaments, the scenario between accretion rate and fragmentation might be more complicated, as discussed by Clarke et al. (2016).

The formation of large-scale structures that are divided into smaller sub-filaments is also seen in simulations, such as those discussed by Smith et al. (2016). In their hydrodynamic turbulent cloud simulations, the sub-filaments are formed naturally and not as a result of fragmentation. Thus, core formation seems to have little or no effect on gas kinematics, suggesting that core formation through violent cloud-to-cloud collision is rare, although still possible (Balfour et al. 2015).

Despite the numerous observational constraints, the exact formation and destruction methods of filamentary structures are still not clear. Rivera-Ingraham et al. (2016, 2017) studied a wide range of close-by ($D < 500$ pc) filaments present in the *Herschel* fields. A significant dispersion in the observed filament properties suggests that the evolutionary paths of filaments are highly dependent on the external conditions, such as Galactic location or main physical processes. Rivera-Ingraham et al. (2016, 2017) do not explicitly examine the role of magnetic fields in filament formation. The orientation of magnetic fields has been observed to be linked to column density (Planck Collaboration Int. XXXV 2016; Cox et al. 2016). For instance, in the high latitude cloud L1642 (Malinen et al. 2016), the magnetic field shows a clear transition in direction from aligned to perpendicular to the cloud elongation at a column density of approximately $N_H = 1.6 \times 10^{21} \text{ cm}^{-2}$. The findings of Rivera-Ingraham et al. (2016, 2017) suggest that the formation

and evolution of filaments is driven by complex non-equilibrium processes between intrinsic and environmental conditions. Thus, understanding the differences in the primordial local conditions is vital in order to determine the structure and evolution of the clouds as well as star formation therein.

The *Herschel* open time key programme Galactic Cold Cores (GCC) carried out dust continuum emission observations of 116 fields that were selected based on the *Planck* Catalogue of Galactic Cold Clumps (PGCC). The fields were mapped with *Herschel* PACS and SPIRE instruments (Pilbratt et al. 2010; Poglitsch et al. 2010; Griffin et al. 2010) at wavelengths of 100–500 μm . *Herschel* makes it possible to study the *Planck* clumps and filaments and their internal structure in detail (Juvela et al. 2012; Montillaud et al. 2015; Rivera-Ingraham et al. 2015). First results from *Planck* and *Herschel* have been presented in Planck Collaboration XXIII (2011), Planck Collaboration XXII (2011), and in Juvela et al. (2010, 2011, 2012). Montillaud et al. (2015) present an analysis of submillimetre clumps and star formation in these *Herschel* fields.

For this paper, we have selected one of the GCC fields, G82.65-2.00. The filament was identified from *Planck* observations as a region of particularly cold dust emission, with the cold dust concentrated to a filamentary structure (Figs. 1–3). The observations of the field revealed a very prominent, strongly fragmented cloud filament, rich in young stellar objects (YSOs) and starless clumps. In total, 28 YSOs and 24 cores have been found to be associated with the main filament (Montillaud et al. 2015). With the assumed distance of 620 pc (see Sect. 3) the mass of the filament exceeds $1200 M_\odot$. In this paper we utilise line observations of several molecules, including CO isotopologues, HCO^+ , HCN, and CS from the Osaka 1.85 m telescope and the Nobeyama 45 m telescope to add kinematical information to study possible correlations between the gas and dust components.

The morphology of G82.65-2.00 is filamentary, although fragmented. The strong fragmentation indicates that the filament is more evolved than, for example, the Taurus filaments B211/213 and TMC-1 or Musca. Compared to high-mass star forming infrared dark clouds, it is less massive and isolated being located at $b = -2^\circ$, suffering much less from confusion by other clouds on the same line-of-sight. The fragmented nature of the region, and the clear evidence of core formation, may infer that the substructures are dissolving back to the interstellar medium (ISM). If the filament is truly dissolving it would be a prime candidate to study the evolutionary stage of a “debris filament”, an evolutionary stage of filaments after star formation but before the structures have dissipated to the ISM.

The content of this paper is as follows. In Sect. 2, we give an overview of our observations. In Sect. 3, we present our main results, and in Sect. 4, we provide some discussion concerning our main results. In Sect. 5, we give our concluding remarks.

2. Observations

2.1. Dust continuum observations

The *Herschel* observations are discussed in detail by Juvela et al. (2012) and Montillaud et al. (2015). The SPIRE observations at 250 μm , 350 μm , and 500 μm were reduced with the *Herschel* Interactive Processing Environment (HIPE) v.12.0, using the official pipeline with the iterative destriper and the extended emission calibration options. The maps were produced with the naive map-making routine. The accuracy of the absolute calibration of the SPIRE observations is expected to be better than 7%. The

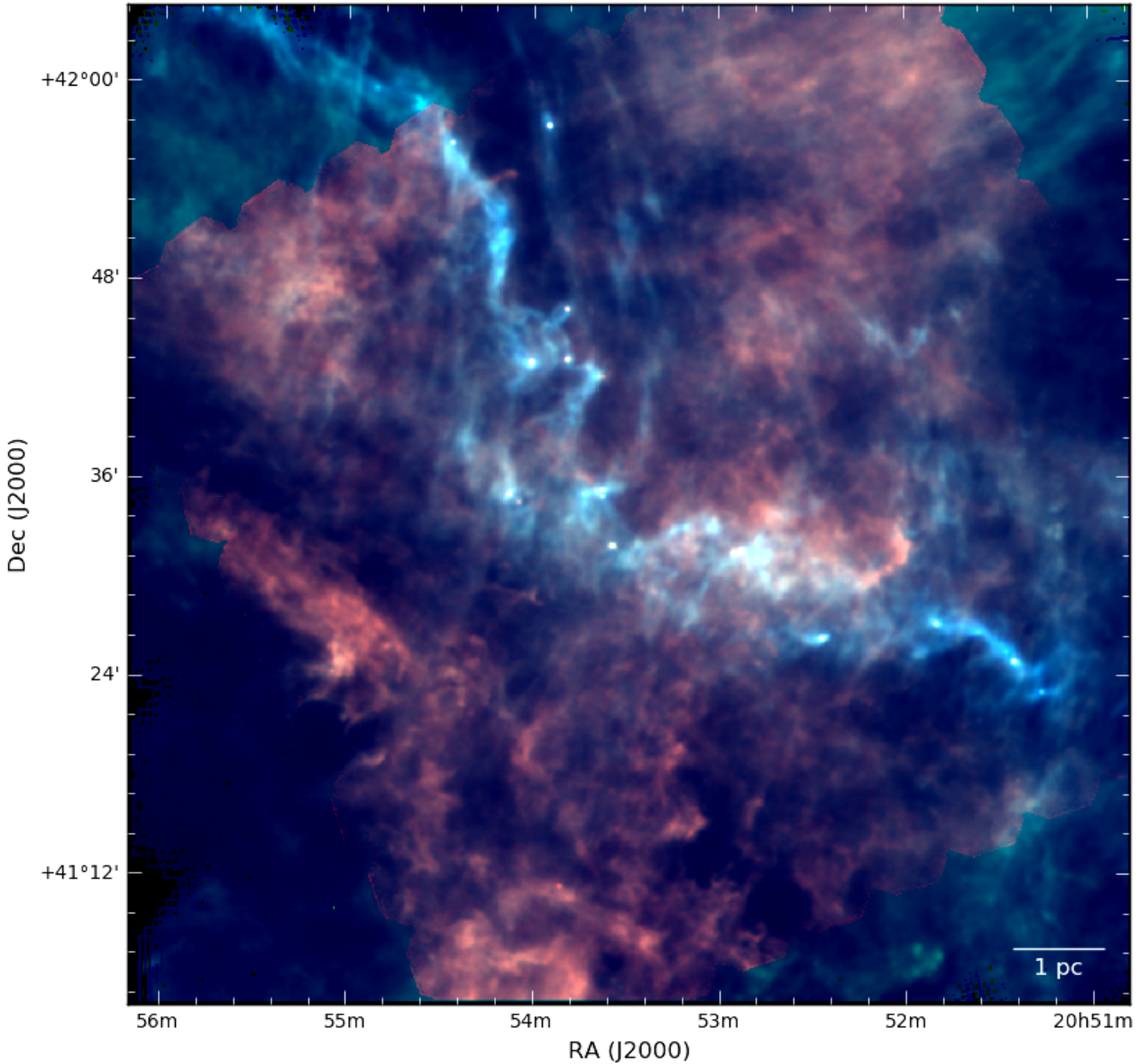


Fig. 1. RGB image of the field G82.65-2.00. The colours correspond to *Herschel* channels at 160 μm (red), 250 μm (green), and 350 μm (blue).

PACS data at 160 μm were processed with HIPE v.12.0 up to Level 1 and the final map was produced with Scanamorphos v.23 (Roussel 2013). For PACS, when combined with SPIRE data, we adopt an error estimate of 10%¹. The value is consistent with the differences between PACS and *Spitzer* MIPS measurements of extended emission. In order of increasing wavelength, the resolution is approximately 12'', 18'', 25'', and 37'' for the four bands. However, the total size of the filament is larger than our *Herschel* maps. The densest ridge of the filament is visible in absorption, although faintly, even in the Wide-field Infrared Survey Explorer (WISE; Wright et al. 2010) 12 μm map in Fig. 2, extending over ~ 25 pc in total.

2.2. Large-scale CO observations with the Osaka 1.85 m telescope

In order to study the velocity field as well as the molecular distributions, we observed an area covering $\sim 1.25^\circ \times 1.25^\circ$,

¹ <http://Herschel.esac.esa.int/twiki/bin/view/Public/PacsCalibrationWeb>

corresponding to the white rectangle in Fig. 2, with the $^{12}\text{CO}(J = 2-1)$, $^{13}\text{CO}(J = 2-1)$, and $\text{C}^{18}\text{O}(J = 2-1)$ emission lines using the 1.85 m Osaka telescope (Onishi et al. 2013). The observations were carried out in September 2013 in the same way as for the observations of Monkey Head nebulae reported by Shimoikura et al. (2013) and using the on-the-fly (OTF) mapping technique.

The half power beam width (HPBW) of the telescope is 2.7' at 230 GHz, the rest frequency of the $^{12}\text{CO}(J = 2-1)$ emission line. We used a digital spectrometer providing the ~ 54 kHz frequency resolution which corresponds to ~ 0.07 km s⁻¹ velocity resolution at 230 GHz. The velocity range of about ± 100 km s⁻¹ was covered by the observations. Noise level of the resulting data is ~ 0.6 K in units of T_{mb} . For more details of the telescope and the observational methods, see papers by Onishi et al. (2013) and Shimoikura et al. (2013).

The C^{18}O data suffers from a poor signal-to-noise ratio (S/N), because the emission line is weak, and thus we use only ^{12}CO and ^{13}CO data in our analyses. We display velocity channel maps of the ^{12}CO and ^{13}CO emission lines in Figs. A.1 and A.2.

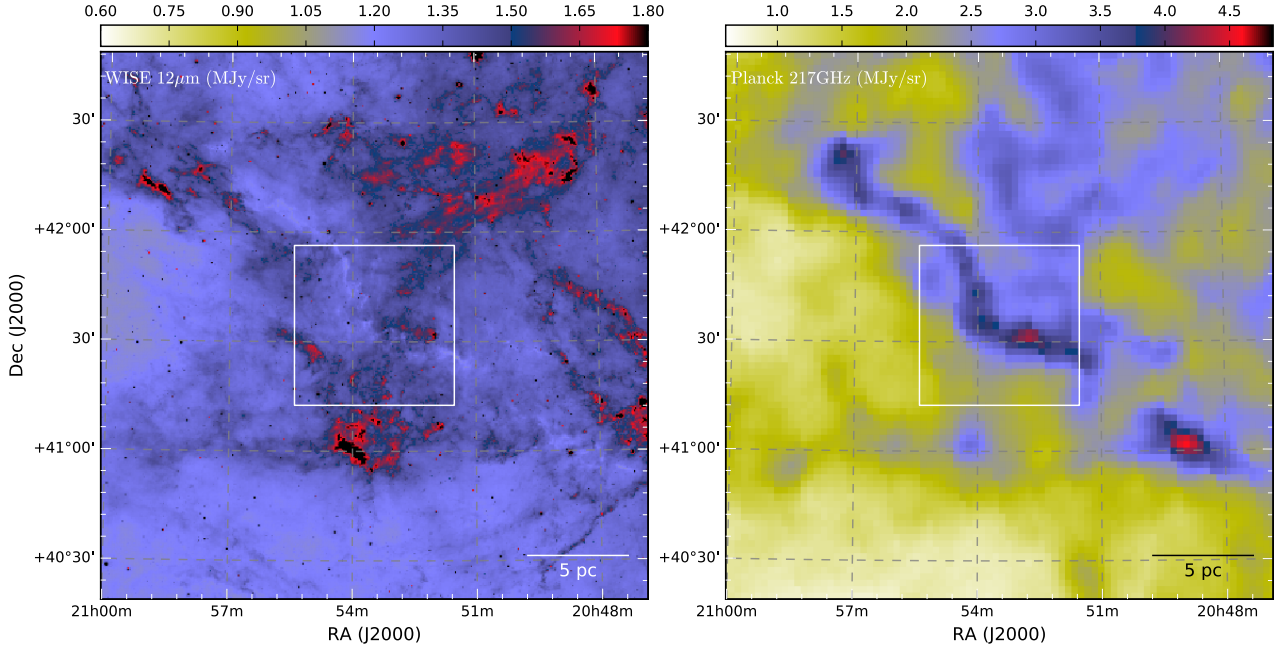


Fig. 2. WISE 12 μm map (left) and *Planck* 217 GHz map (right). The *Planck* map has been converted to MJy/sr using the conversion factors given in *Planck Collaboration IX (2014)*. The white rectangle shows the extent of the maps in Fig. 3.

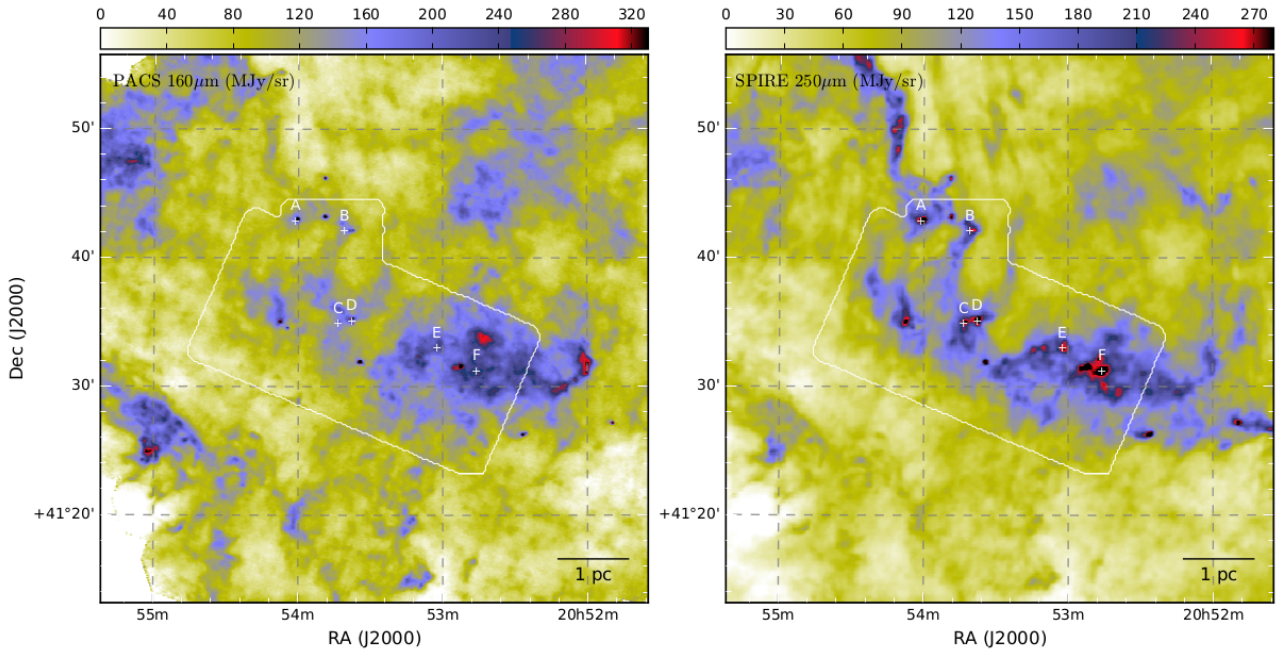


Fig. 3. Surface brightness observed by *Herschel*: PACS 160 μm map (left), and SPIRE 250 μm map (right). The labels from A to F and the white crosses show the locations of cold clumps analysed in Sect. 3. The white contour shows the area covered by our line observations from the Nobeyama 45 m telescope.

2.3. Molecular line observations with the Nobeyama 45 m telescope

Further spectral line observations were carried out with the 45 m telescope at the Nobeyama Radio Observatory (NRO) for 15 days in April and May 2014. We used the waveguide-type dual-polarisation sideband-separating SIS receiver TZ (*Nakajima et al. 2013*) with the digital spectrometer SAM45. The set-up provides 4096 channels with a ~ 8 kHz frequency resolution covering an ~ 32 MHz bandwidth. The corresponding

velocity resolution and velocity coverage at 110 GHz are ~ 0.025 km s $^{-1}$ and ~ 100 km s $^{-1}$, respectively. The beam size of the telescope at 110 GHz is $\sim 14.2''$ (HPBW).

We observed the $^{12}\text{CO}(J = 1-0)$, $^{13}\text{CO}(J = 1-0)$, $\text{C}^{18}\text{O}(J = 1-0)$, $\text{CS}(J = 2-1)$, and $\text{HCO}^+(J = 1-0)$ emission lines. The combination of the TZ receiver and the SAM45 spectrometers, however, enables us to observe some additional molecular lines simultaneously (up to sixteen lines) with 8 to 16 GHz frequency separation. In total, we observed twenty molecular lines as summarised in Table 1.

Table 1. Observed molecular lines from Nobeyama.

Molecule	Transition	Rest frequency (GHz)	$\Delta T_{\text{mb}}^{(1)}$ (K)	Detection
^{12}CO	$J = 1-0$	115.271202	0.53	Yes
CN	$J = 1-0; J = 3/2-1/2; F = 5/2-3/2$	113.490982	0.40	No
CCS	$N; J = 9; 8-8; 7$	113.410204	0.41	No
C^{17}O	$J = 1-0$	112.358988	0.35	No
CH_3CN	$J = 6(2) - 5(2); F = 7-6$	110.375052	0.21	No
^{13}CO	$J = 1-0$	110.201353	0.20	Yes
NH_2D	$J = 1(1; 1)0-1(0; 1)0+$	110.153599	0.22	No
C^{18}O	$J = 1-0$	109.782173	0.19	Yes
H_2CS	$J = 3(1; 3) - 2(1; 2)$	101.477885	0.21	No
HC_3N	$J = 11-10$	100.076385	0.16	Yes
SO	$N; J = 5; 4-4; 4$	100.029565	0.16 (0.15)	No
SO	$N; J = 2; 3-1; 2$	99.299905	0.21 (0.15)	Yes
CS	$J = 2-1$	97.980953	0.13 (0.13)	Yes
CH_3OH	$J = 2(0; 2) - 1(0; 1)A ++$	96.741377	0.16	Yes
C^{34}S	$J = 2-1$	96.412961	0.15	No
CH_3OH	$J = 2(1; 2) - 1(1; 1)A ++$	95.914310	0.14	No
HCO^+	$J = 1-0$	89.188526	0.19 (0.13)	Yes
HCN	$J = 1-0; F = 2-1$	88.6318473	0.18 (0.13)	Yes
SiO	$J = 2-1; v = 0$	86.846995	0.22 (0.16)	No
H^{13}CO^+	$J = 1-0$	86.754288	0.18 (0.14)	No

Notes. Values in parentheses are those for the additional observations to map a small core, clump B in Fig. 6, located beside the main filament.

⁽¹⁾ Typical one sigma noise level (rms) at the 0.2 km s^{-1} velocity resolution.

We mapped an area of $12' \times 24'$ around the main filament using the OTF technique (Sawada et al. 2008), and calibrated the spectral data with the standard chopper-wheel method (Kutner & Ulich 1981). The reference position for the OTF observations (i.e. the emission-free OFF positions) was chosen to be $(\alpha_{J2000}, \delta_{J2000}) = (20^{\text{h}}57^{\text{m}}05.04^{\text{s}}, 39^{\circ}19'59.4'')$. During the observations, we discovered a well-defined small core located just outside of the $12' \times 24'$ area covering the main filament. Therefore, we carried out an additional OTF mapping covering a $\sim 2' \times 2'$ area around the small core in certain molecular lines (see Table 1). The extent of the Nobeyama observations is shown in Fig. 3.

We used the reduction software package NOSTAR available at NRO to subtract a linear baseline from the raw data and to re-sample the spectral data onto a $10''$ grid aligned with the equatorial coordinates. A varying main beam efficiency correction, in the range of $\eta_{\text{mb}} = 29.2\text{--}46.6\%$ depending on the observed frequencies, was applied to scale the data to units of T_{mb} .

The frequency channels were re-sampled onto a 0.025 km s^{-1} velocity grid, and then smoothed to $0.2\text{--}0.4 \text{ km s}^{-1}$ resolution to reduce the noise. Noise levels are in the range of $\Delta T_{\text{mb}} = 0.1\text{--}0.5 \text{ K}$ for a velocity resolution of 0.2 km s^{-1} , as summarised in Table 1. Among the twenty molecular lines we observed, nine of them were detected with S/N greater than 3. These lines are labelled “Yes” in the last column of the table. We display velocity channel maps of the nine emission lines in a series of figures (Figs. B.1 to B.9).

The system noise temperature including the atmospheric attenuation varied between 140 and 250 K depending on the observed frequency. Pointing accuracy was better than $\sim 10''$, and was checked by observing a SiO maser, T-Cep, at 43 GHz and a H_2O maser, W75N, at 22 GHz every 1–2 h.

3. Results

3.1. Distance estimate

A distance estimate is crucial to derive the mass of molecular clouds and to investigate their kinematics. However, the distance to G82.65-2.00 has not been accurately determined to date. Montillaud et al. (2015) estimated the distance to be $1 \pm 0.5 \text{ kpc}$ based on several methods but the scatter is large depending on the method. CO observations of Dame et al. (2001) show that the field around the filament consists of a single component at 3.47 km s^{-1} , in agreement with our results showing the main line emission component in the range of $[3, 5] \text{ km s}^{-1}$ (see Sect. 3.2). In general, distance can be derived from spectroscopic observations, assuming that the galactic gas is in circular rotation. Using the Galactic rotation curve by Reid et al. (2009, 2014), Montillaud et al. (2015) derive two kinematics distances of 3.30 ± 0.14 and $1.08 \pm 0.76 \text{ kpc}$ assuming a tangential peculiar velocity of $V_0 = 0$ and $V_0 = -15 \text{ km s}^{-1}$, respectively. This latter value corresponds to the average orbital velocity reported by Reid et al. (2009, 2014) for massive star forming regions. However it remains unclear whether or not $V_0 = -15 \text{ km s}^{-1}$ is correct for G82.65-2.00.

Other studies of clouds that seem to be connected with G82.65-2.00 have reported similar distance estimates. Dobashi et al. (1994) proposed a distance of 800 pc for their cloud 16 ($l = 82.93^\circ$, $b = -2.03^\circ$), while Bally & Scoville (1980) adopted a distance of 1 kpc for the Pelican Nebula ($l = 84.63^\circ$, $b = +0.10^\circ$). A recent work by Cersosimo et al. (2007) shows that several structures are seen in the region of the Pelican nebula, corresponding to different velocity components. The closest cloud to G82.65-2.00 in their study ($l = 84^\circ$, $b = -1.7^\circ$), has similar velocity of 2.2 km s^{-1} for which Cersosimo et al. (2007) derived a distance of $0.5\text{--}0.7 \text{ kpc}$, and suggested greater distances (1.7 kpc, 2.7 kpc, and 3.3 kpc) for the other structures.

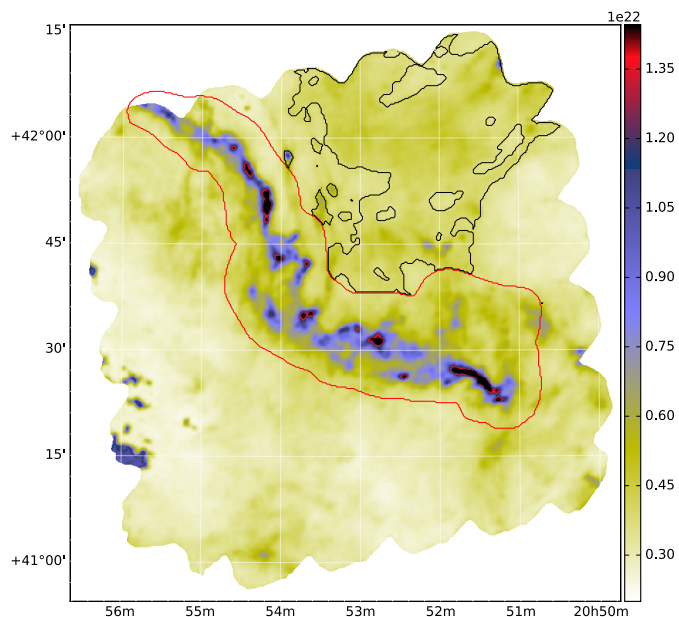


Fig. 4. Areas used by the extinction method to derive distance estimates. The area inside the red contour is used to estimate the distance of the main filament, and the area inside the black contour is used to estimate the distance of the secondary component.

Because of the large uncertainty in the kinematic distance estimate, [Montillaud et al. \(2015\)](#) derived an independent estimate using extinction method. By comparing observed stellar colours from the Two Micron All Sky Survey ([Skrutskie et al. 2006](#), 2MASS) point-source catalogue with the predictions of the Besançon Galaxy model ([Robin et al. 2003, 2012](#)), they derived the most probable three dimensional extinction distribution along the line-of-sight. The method is described by [Marshall et al. \(2006, 2009\)](#). [Montillaud et al. \(2015\)](#) applied the method to two circular areas with 5' and 10' in diameter, centred on the brightest position of the *Herschel* 250 μm map, which corresponds to clump F in Fig. 3. The estimated distances were 0.42 kpc for the small aperture and 0.98 kpc for the large aperture. The difference between the two values is included in the formal uncertainties provided by the method.

Considering the large uncertainties between the different methods, we decided to obtain a more reliable estimate for the distance using an updated version of the extinction method described above ([Marshall et al. 2015](#), and in prep.). [Montillaud et al. \(2015\)](#) used only a narrow area around clump F, and thus the number of stars that can be used in the method was limited, which induced the large uncertainty in the distance estimate. We have increased the area in the extinction method to cover the entire cold filament as indicated by the red contour in Fig. 4. Looking at the *Herschel* data and the WISE 12 μm map (Figs. 1 and 2), it is likely that there is a second, warm and more extended filamentary structure that runs through the whole field from south-east to north-east. The structure is visible in Fig. 1 in red. We have used the extinction method to derive a distance estimate for this warmer filament using the area indicated by the black contour in Fig. 4.

The computations were performed assuming either one or two clouds on the line-of-sight. For the cold filament, the highest posterior probability is obtained with one cloud at 620_{-42}^{+31} pc. For the warm filament, the results varied depending on the area to apply the method. With a smaller area, as seen in Fig. 4, the highest probability is reached with one component at 300 pc.

However, using a wider area the highest probability is obtained by two clouds, with one cloud at 661_{-55}^{+69} pc and a second cloud at $7.3_{-1.6}^{+0.7}$ kpc. A distance of 300 pc is plausible, however, a distance of ~ 7 kpc is unlikely as the field is resolved with *Herschel*.

Thus, for the cold filament, we will adopt the distance 620 pc in this paper. The distance of the warm filament is more uncertain and requires further studies.

3.2. Large-scale kinematics

The Osaka ^{12}CO data, Fig. 5, reveals an intriguing overall kinematic structure, with one filament visible in the range $\sim[-5, -2]$ km s^{-1} and extended emission appearing in the range $\sim[0, 10]$ km s^{-1} . At low radial velocities (see Fig. A.1 for more detailed velocity channel maps), in the range $\sim[0, 4]$ km s^{-1} , the eastern edge of the extended emission seems to trace the cold filament well. At velocities higher than 7 km s^{-1} , there is a coherent velocity structure perpendicular to the central part of the cold filament. The large-scale Osaka ^{12}CO data would thus imply that the cold filament seen in the *Herschel* data is located between three, and possibly four, different velocity structures.

The three compact sources seen in the lower left corner of the channel maps, Fig. A.1, in the range $\sim[-6, 1]$ km s^{-1} are probably related to the Infrared Astronomical Satellite (IRAS) point source 20529+4114. Although the fainter object in the range $\sim[-6, -4.5]$ km s^{-1} is within a few arc seconds of the IRAS source, the peak CO emission from the three sources is offset by $\sim 3'$.

In the Osaka ^{13}CO maps (Fig. A.2) the only prominent feature is a filament that matches the cold filament seen in the *Herschel* data. The emission reaches maximum at ~ 4 km s^{-1} and concentrates around the main filament. At lower radial velocities, the emission is found in the north-eastern and south-western part of the filament.

For the Nobeyama line observations, we only show channel maps in the velocity range $\sim[0.0, 9.5]$ in Fig. B.1. Although ^{12}CO was also detected in the range of -4.7 to -2.5 km s^{-1} , we have not included this velocity range in our analyses as it is distributed rather randomly, inferring that there is no relation to the main filament. The channel maps of the Nobeyama observations are shown in Appendix B.

The Nobeyama ^{12}CO emission in Fig. B.1 follows the same pattern as the Osaka data. The peak of the emission moves from the north-east towards south-west, tracing well the dust emission in each velocity interval. Some well defined substructures with different velocities can be identified, for example, in the range $\sim[-0.4, 1.5]$ km s^{-1} and in the range $\sim[7, 9]$ km s^{-1} .

As expected, and as seen in the data from Osaka, the ^{13}CO line traces better the shape of the cold filament visible in dust emission in Fig. 3. Some individual substructures can be identified from the contour maps. In the range of $\sim[4, 6]$ km s^{-1} , the ^{13}CO line clearly traces the striation S1, marked with a cyan ellipse in Fig. 6, and connects the striation to the dense clumps C and D.

The Nobeyama C^{18}O data (Fig. B.3) trace the densest region of the cold filament but the line is also seen on the striation S1 (see Fig. 6) and, as with the Nobeyama ^{13}CO , the line emission extends to clumps C and D.

In order to study the gas dynamics, we computed pixel-by-pixel Gaussian fits for the Nobeyama ^{12}CO , ^{13}CO , and C^{18}O data. To reduce the noise level in the resulting images, the data were averaged over $30'' \times 30''$ corresponding to 3×3 map pixels. The Gaussian fits return the three line parameters: the peak

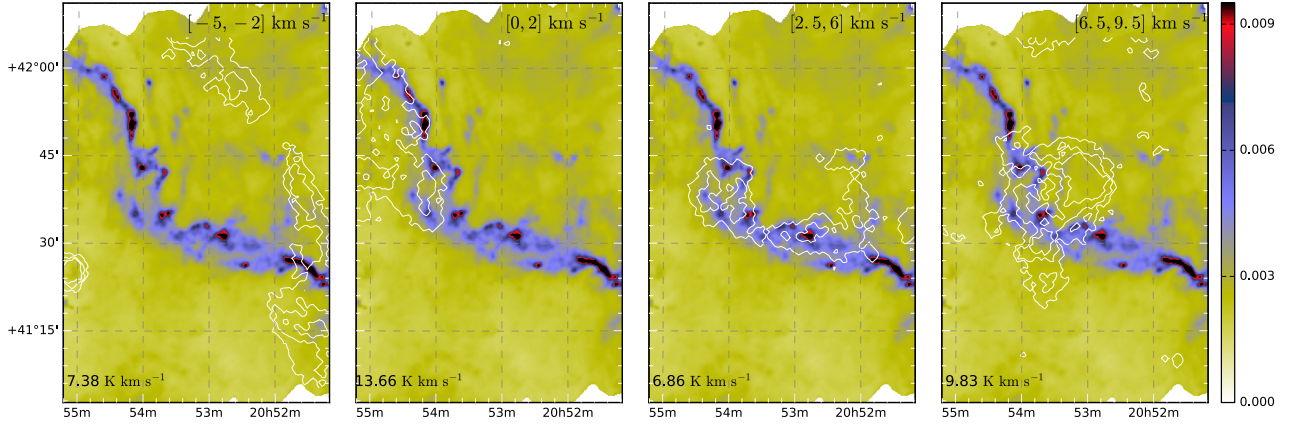


Fig. 5. Different velocity structures identified from the Osaka observations. The underlying colour map is the τ_{250} map. The contours have been computed from the Osaka ^{12}CO data and show 50%, 70%, and 90% of the peak emission in the velocity interval noted in the upper right corner. The peak emission in each velocity interval is denoted in the lower left corner.

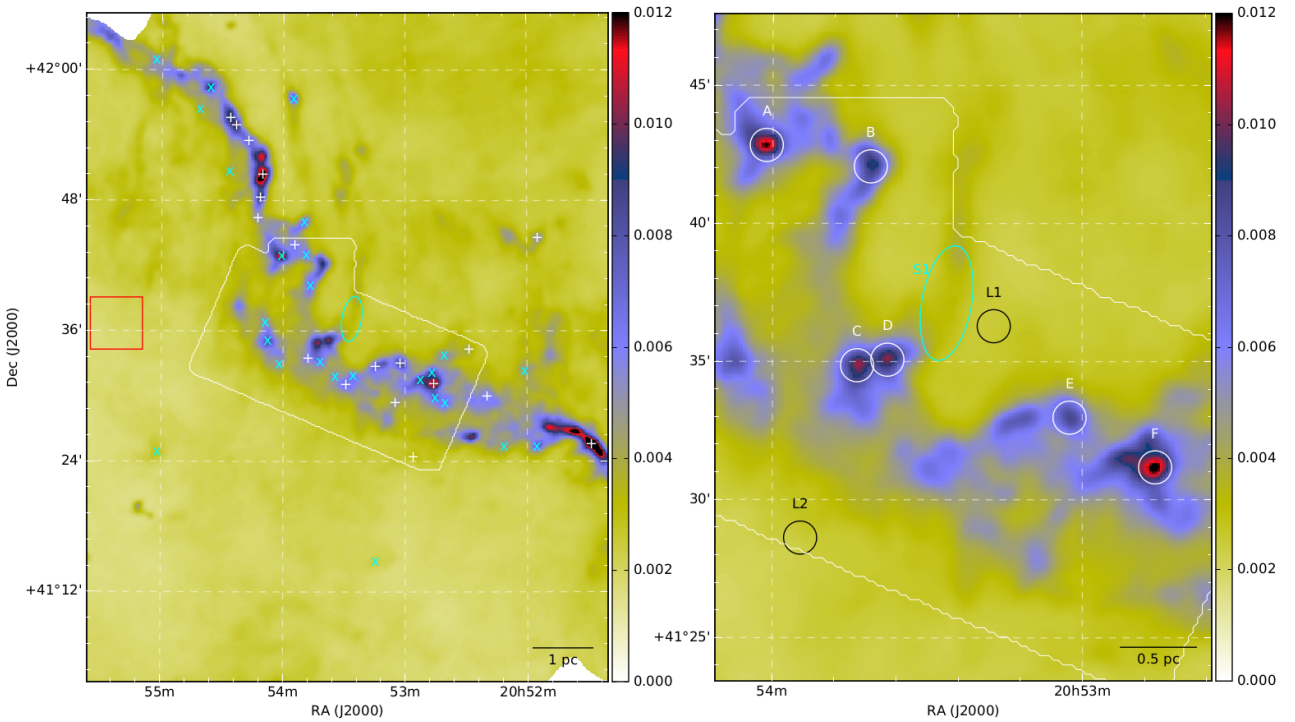


Fig. 6. *Left panel:* optical depth map of 250 μm derived from the *Herschel* observations, the white plus signs indicate starless cores and the cyan crosses indicate protostellar source candidates identified by [Montillaud et al. \(2015\)](#); and *right panel:* clumps (labelled A to F) chosen for the spectral analysis (Figs. 14 and 15). The labels L1 and L2 indicate the positions where we have extracted spectra to analyse the warm filament (Fig. D.1). The white outline in the panels corresponds to the area of the Nobeyama observations. The cyan ellipse marks the selected striation S1, and the red rectangle marks the area used for the background subtraction as discussed in Sect. 3.4.

brightness temperature T_{mb} , the radial velocity V_0 , and the line width ΔV measured at full width at half maximum (FWHM). The results for V_0 and ΔV are shown in the second and third columns in Fig. 7. Because of the different velocity components, the shape of the ^{12}CO line is complex, and thus, a single Gaussian fit may not describe the properties of the line well. However, results described below do not change even when we fit the spectra with two Gaussian components.

Figure 7a shows the observed total integrated intensity of ^{12}CO . The emission reaches maximum on the eastern side of the image, and traces the same region as the velocity component seen in the range $\sim[-0.5, 2]$ km s^{-1} in Fig. B.1. The overlaid contours show the distinct emission component in the range

$\sim[-4, -1]$ km s^{-1} that is not used in the analysis. The velocity map in Fig. 7b shows that the peak velocities are in the range $\sim[3, 5]$ km s^{-1} , with a notable exception in the central part of the panel. The higher-velocity component corresponds to the velocity structure seen in the range $\sim[7, 9]$ km s^{-1} in Fig. B.1. The two velocity components are also clearly distinguishable in the line width map in Fig. 7c, but the larger line width also corresponds to the eastern part of the filament, where the integrated emission peaks.

The velocity and line width maps of ^{13}CO and C^{18}O (Figs. 7f and i) show the main emission in the range $\sim[3, 5]$ km s^{-1} . However the features seen in the ^{12}CO maps, the clearly distinguishable higher-velocity component and two components in the line

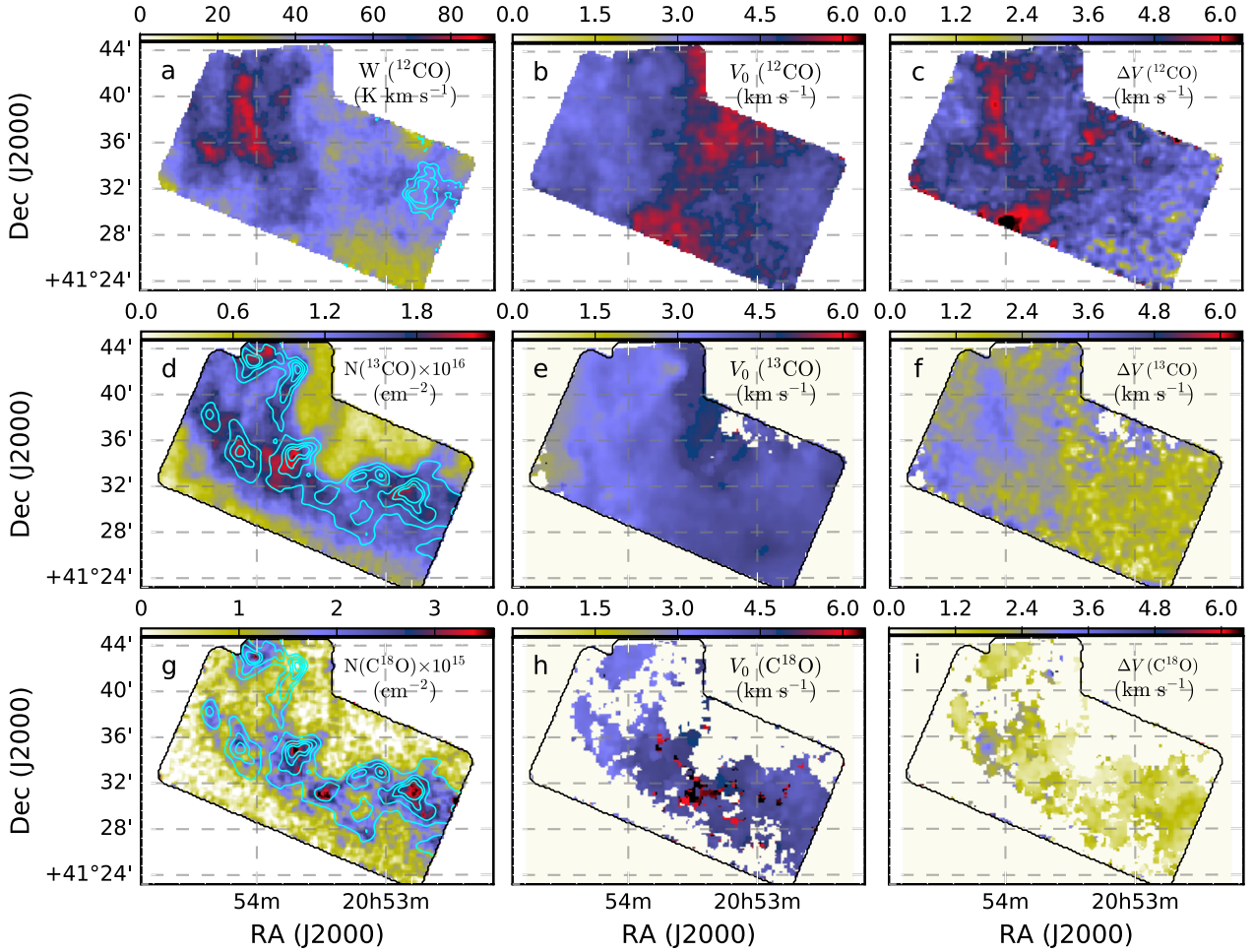


Fig. 7. First row: total integrated intensity (K km s^{-1}) of ^{12}CO , velocity and line width at FWHM (km s^{-1}) obtained from the Gaussian fit of the ^{12}CO line. Second and third rows: total column densities and best-fit values for velocity and line width obtained from the Gaussian fit of the ^{13}CO and C^{18}O lines, respectively. The column densities are derived with a constant excitation temperature of 15 K. The contour in the $W(^{12}\text{CO})$ map shows the observed ^{12}CO emission in the range of $[-4, -1] \text{ km s}^{-1}$ which was not included in our analysis. The lowest contour is at 0.7 K km s^{-1} and the contours increase with 0.7 K km s^{-1} steps. The contours in the ^{13}CO and C^{18}O column density maps correspond to 45%, 60%, 75%, and 90% of the maximum value of τ_{250} . We have masked regions with $S/N < 3$ in the velocity and line width maps of ^{13}CO and C^{18}O .

width map, are not seen, because ^{13}CO and C^{18}O are not detected at velocities higher than $\sim 5.5 \text{ km s}^{-1}$. The white areas in the maps were masked out because of low S/N .

3.3. Column density estimates from line observations and dust emission

Shown in Fig. 8, is a dust temperature map derived from the *Herschel* observations. The temperature was estimated by fitting the spectral energy distributions (SEDs) with modified blackbody curves with a constant opacity spectral index of $\beta = 1.8$, which is the average value in molecular clouds (Planck Collaboration XXII 2011; Planck Collaboration XXIII 2011; Planck Collaboration XXV 2011; Juvela et al. 2015a,b). The cold filament, seen in blue in Fig. 8, contains substructures of a few times 0.1 pc , including several cold clumps. In the *Herschel* images, one can identify a number of fainter striations, some of which are indicated by white arrows in Fig. 8.

We use the SPIRE observations at $250 \mu\text{m}$, $350 \mu\text{m}$, and $500 \mu\text{m}$, to estimate the dust optical depth at $250 \mu\text{m}$. The surface-brightness maps were convolved to a common resolution of $40''$, and colour temperatures were calculated by fitting the

SEDs with modified blackbody curves with a constant opacity spectral index of $\beta = 1.8$. The $250 \mu\text{m}$ optical depth was obtained from

$$\tau_{250} = \frac{I_\nu(250 \mu\text{m})}{B_\nu(T_c)}, \quad (1)$$

using the fitted $250 \mu\text{m}$ intensity $I_\nu(250 \mu\text{m})$ and the colour temperature T_c . The data were weighted according to 7% error estimates for SPIRE surface-brightness measurements.

The *Herschel* observations are used to derive an estimate of the total hydrogen column density. The optical depth τ_ν at frequency ν can be expressed as

$$\tau_\nu = \kappa_\nu N(\text{H}_2) \hat{\mu}, \quad (2)$$

where $\hat{\mu}$ is the total gas mass per H_2 molecule, 2.8 amu , and $N(\text{H}_2)$ is the hydrogen column density. We assume a dust opacity of $\kappa_\nu = 0.1(\nu/1000 \text{ GHz})^\beta \text{ cm}^2 \text{ g}^{-1}$ (Beckwith et al. 1990). The resulting optical depth map is shown in Fig. 6.

The Nobeyama C^{18}O data are used to derive an independent estimate for the hydrogen column density. The column density

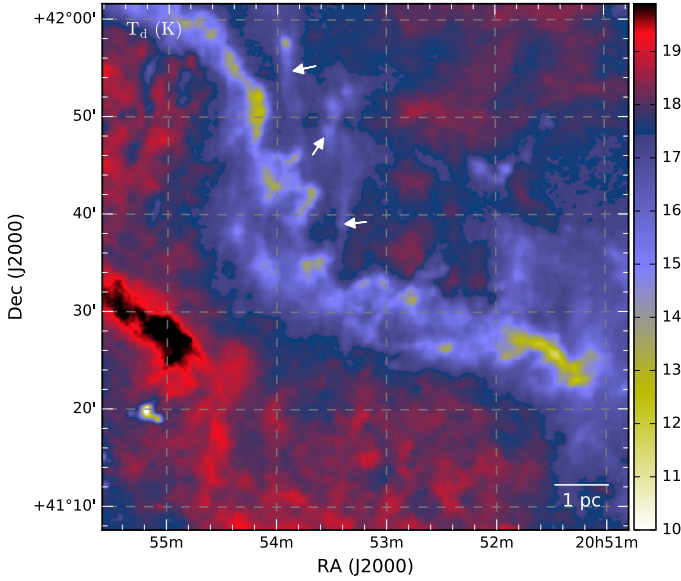


Fig. 8. Dust temperature map derived from the *Herschel* observations. The map is truncated at 19.9 K to show the fainter structures. The white arrows indicate structures identified as striations.

N of C^{18}O is derived following [Garden et al. \(1991\)](#) as

$$N = \frac{3k_b}{8\pi^3 B \mu_r^2} \frac{\exp(hBJ(J+1)/k_b T_{\text{ex}})}{J+1} \times \frac{T_{\text{ex}} + hB/3k_b}{1 - \exp(-h\nu/k_b T_{\text{ex}})} \int \tau_{\nu} dV, \quad (3)$$

where B and μ_r are the rotational constant and the permanent dipole moment of the molecule, J is the rotational quantum number of the lower energy level ($J = 0$), h is the Planck constant, k_b is the Boltzmann constant, and T_{ex} is the excitation temperature.

For the integral of the optical depth, we followed [Buckle et al. \(2010\)](#), with the exception that we do not use an averaged optical depth over the whole region, but rather compute the optical depth τ separately for each pixel and sum over the velocity range where we have detected C^{18}O ,

$$\int \tau_{\nu} dV = \frac{1}{J(T_{\text{ex}}) - J(T_{\text{bg}})} \frac{\tau}{1 - e^{-\tau}} \int T_b dV, \quad (4)$$

where T_{bg} is the cosmic background temperature 2.73 K, T_b is the observed brightness temperature of a given pixel, and $J(T)$ is the temperature equivalent to the temperature T in the Rayleigh-Jeans law (e.g. [Garden et al. 1991](#)). The excitation temperature estimated from the peak brightness temperature of the Nobeyama ^{12}CO data varies over the observed region (up to ~ 40 K) with an average value of ~ 15 K. In this paper, we assume a flat excitation temperature of $T_{\text{ex}} = 15$ K. We note that the estimates of the molecular column densities are not very sensitive to T_{ex} , and assuming $T_{\text{ex}} = 40$ K produces a higher column density by only a factor of ~ 2 .

We use the column density $N(\text{C}^{18}\text{O})$ to derive an estimate for the total hydrogen column density $N(\text{H}_2)$. We assume a relative abundance of 1.0×10^4 between H_2 and ^{12}CO and a relative abundance of 4.9×10^2 between ^{16}O and ^{18}O ([Garden et al. 1991](#)). Thus, the resulting hydrogen column density is

$$N(\text{H}_2) = 4.9 \times 10^6 N(\text{C}^{18}\text{O}). \quad (5)$$

For comparison, using the column density $N(^{13}\text{CO})$ results in hydrogen column densities within $\pm 10\%$ of the values derived using $N(\text{C}^{18}\text{O})$.

The column densities of both ^{13}CO and C^{18}O (Figs. 7d and g) trace well the maxima of the τ_{250} distribution. The total hydrogen column density derived from the line observations is $1.0\text{--}2.3 \times 10^{22} \text{ cm}^{-2}$, in very good agreement with the values derived from the *Herschel* observations, $0.5\text{--}2.5 \times 10^{22} \text{ cm}^{-2}$.

3.3.1. Correlation between gas and dust

Figure 9 shows the two-dimensional histograms of the pixel-to-pixel correlation between τ_{250} and the column densities of ^{13}CO and C^{18}O derived from the Nobeyama observations. The column density maps from the molecular line observations have been smoothed to the same resolution, $40''$, as the τ_{250} map. It is evident that the relation between τ_{250} and $N(^{13}\text{CO})$ is not linear. We use a fit of the form,

$$N = N_0(1 - \exp[-c(\tau - \tau_{\text{CO}})]), \quad (6)$$

where N_0 is the column density of the CO species at saturation, τ_{CO} is the minimum dust optical depth for CO emission, and c is a scaling factor. The results of the fits are shown in Fig. 9 and in Table 2.

The ^{13}CO column density starts to rise rapidly after the optical depth reaches a threshold value of 2×10^{-3} , and starts to saturate when the optical depth reaches a value of 5×10^{-3} . On the other hand, the C^{18}O emission starts to rise rapidly at a slightly higher threshold value of $\tau_{250} \approx 2.5 \times 10^{-3}$. Also, the correlation between $N(\text{C}^{18}\text{O})$ and τ_{250} is more linear below $\tau_{250} \approx 5.5 \times 10^{-3}$. Furthermore, the column density of C^{18}O does not reach complete saturation.

3.3.2. Position-velocity diagrams

To study velocity gradients along the cold filament and possible signs of mass accretion, we have drawn position-velocity (PV) diagrams along the three cuts shown in Fig. 10. We used the Osaka ^{12}CO data to make two PV diagrams, from position 3 to 4 and from position 5 to 6. The diagrams are shown in Fig. 11. The two segments of the cold filament are uniform, although, there are some substructures close to positions 4 and 5. However, due to low S/N, it is not clear whether the velocity components are separate components, possibly at different distance on the same line-of-sight, or if the velocity components are a result of hierarchical fragmentation as discussed by [Hacar et al. \(2013\)](#).

The Nobeyama ^{13}CO and C^{18}O data indicate that the striation S1 is connected to the main filament. A PV diagram drawn along the striation from position 1 to position 2, Fig. 11, shows a single velocity component between 4.0 and 5.5 km s^{-1} . Thus, the striation is very likely connected kinematically to clumps C and D (see Fig. 6).

3.4. Mass estimates for clumps and substructures

We have identified clumps as dense regions inside the area covered by the Nobeyama observations where the total hydrogen column density derived from dust is greater than $1.5 \times 10^{22} \text{ cm}^{-2}$. In total, we have identified six clumps as shown in Fig. 12 and in the right panel of Fig. 6. The mass of a clump can be written as

$$M_c = N(\text{H}_2)\pi R^2 \hat{\mu}, \quad (7)$$

where R is the radius of the clump, defined as half of the FWHM measured from the *Herschel* column density map. The average

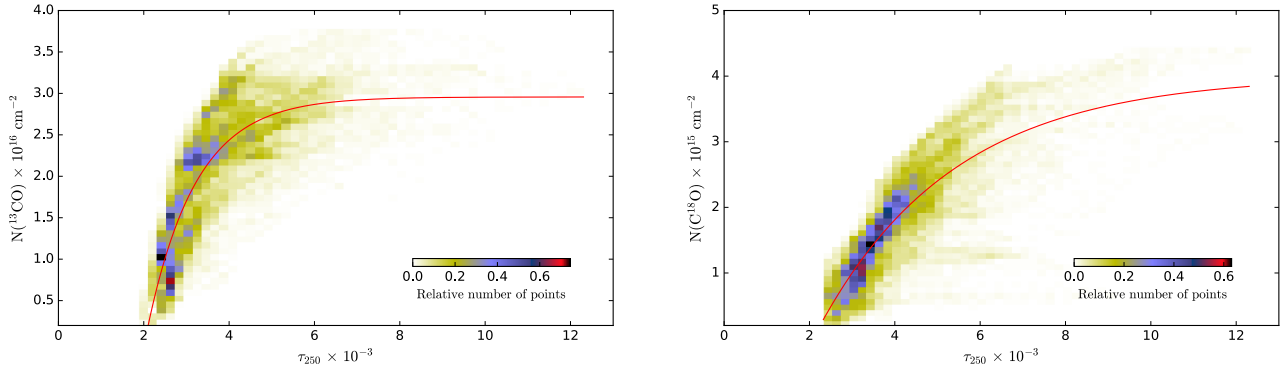


Fig. 9. Pixel-to-pixel correlation of τ_{250} and column density of ^{13}CO and C^{18}O derived from the Nobeyama observations. The red curve is the best fit of Eq. (6). The colour map corresponds to the density of data points.

Table 2. Parameters of the fits to the ^{13}CO and C^{18}O column density as a function of τ_{250} .

Molecule	N_0	τ_{CO}	c
^{13}CO	2.96 ± 0.07	0.88 ± 0.09	2.02 ± 0.8
C^{18}O	4.01 ± 0.3	0.31 ± 0.05	2.09 ± 0.1

Notes. Best-fit values from Eq. (6) to data presented in Fig. 9.

molecular density n of a clump is calculated using the average hydrogen column density \bar{N} inside the radius as

$$n = \frac{\bar{N}(\text{H}_2)}{2R}. \quad (8)$$

To estimate the virial mass, we follow MacLaren et al. (1988)

$$M_{\text{virial}} = \frac{k\sigma^2 R}{G}, \quad (9)$$

where k depends on the assumed density distribution, and G is the gravitational constant. We have chosen a density distribution corresponding to $\rho(r) \propto r^{-1.5}$, and thus, $k \approx 1.33$. The velocity dispersion, σ , can be computed from the equation

$$\sigma = \sqrt{\frac{k_b T_{\text{kin}}}{\mu} + \frac{\Delta V^2}{8 \ln 2} - \frac{k_b T_{\text{kin}}}{m}}, \quad (10)$$

where ΔV is the line width (at FWHM) of the line used, m is the mass of the molecule used in the computation (^{13}CO), and μ is the mean molecular mass, 2.33 amu assuming 10% helium. We have assumed a constant kinetic temperature of $T_{\text{kin}} = 15$ K. The velocity dispersion can be used to estimate the nonthermal motion that provides support against gravity. If the mass of the cloud is less than the virial mass, it is not gravitationally bound and, without external pressure, will disperse.

The average hydrogen column density, average density, velocity dispersions for ^{13}CO and C^{18}O , estimated mass, and virial mass estimates for the clumps are listed in Table 3. The densities and masses are derived from the continuum observations. For the computations, a distance of 620 pc was assumed. The error estimates were calculated with Monte Carlo, assuming normally distributed errors.

To avoid possible contribution by the background in our analysis, a constant background value of $1.2 \times 10^{21} \text{ cm}^{-2}$ was subtracted from the column density values when analysing dust continuum. The background value was estimated from the

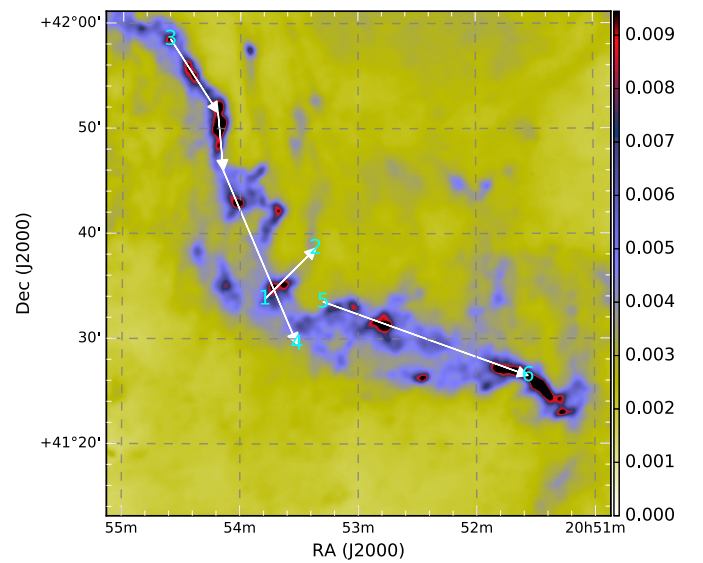


Fig. 10. Cuts used for the position-velocity maps, shown on the τ_{250} map.

Herschel images as the mean value inside the red rectangle in Fig. 6. Thus, we do not take into account any possible contribution from the medium surrounding the clump. However, even for the more extended clumps, mainly clump F, the contribution from surrounding medium is at most $\sim 15\%$.

We have identified some substructures within the filaments by eye, as labelled in Fig. 13. These were first identified from the Osaka ^{13}CO data and compared with the *Herschel* column density maps to identify the structures that are well correlated between gas and dust observations.

The substructure masses were estimated following Eq. (7), with the exception of radius which is defined as $R = \sqrt{S/\pi}$, where S is the surface area of the structure shown by the ellipses in Fig. 13. The surface area is computed as $S = \pi R_{\text{maj}} R_{\text{min}}$, where R_{maj} and R_{min} are the lengths of the semi-major and semi-minor axes of the structure, respectively, as shown in Fig. 13.

In order to investigate the gravitational stability of the substructures, we assume a model of isothermal cylinder and compute the mass per unit length as

$$M_{\text{line}} = \frac{M_f}{2R/\sqrt{\alpha}}, \quad (11)$$

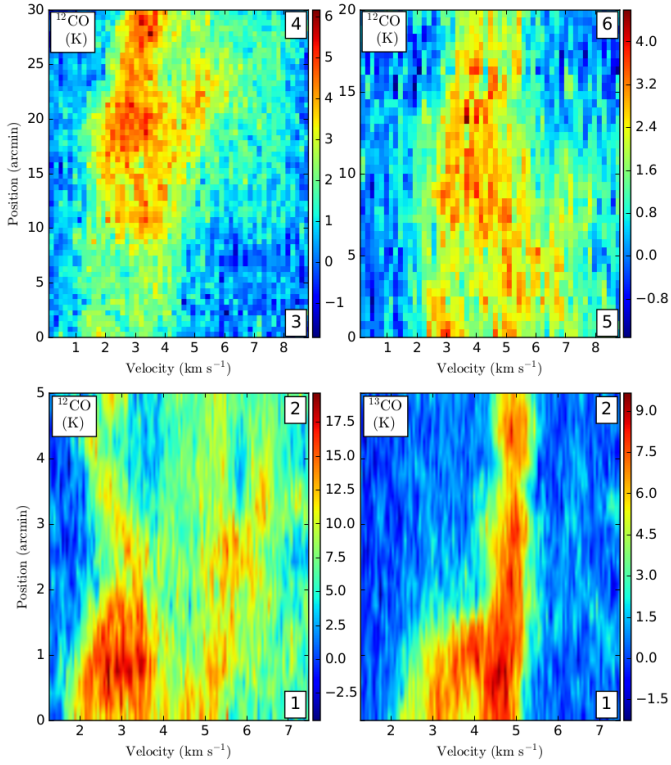


Fig. 11. Position-velocity maps, derived using the Osaka ^{12}CO ($J = 2-1$) data, *upper row*, and the Nobeyama ^{12}CO ($J = 1-0$) and ^{13}CO ($J = 1-0$) data, *lower row*. The numbers in the panels refer to the positions shown in Fig. 10.

where M_f is the mass and $\alpha = R_{\min}/R_{\text{maj}}$ the ellipticity of the substructure. The virial mass per unit length is defined following Fiege & Pudritz (2000),

$$M_{\text{line, vir}} = \frac{2\sigma^2}{G}. \quad (12)$$

Listed in Table 4 are the average hydrogen column density, the average density, the average velocity dispersion, the mass and the M_{line} (Eq. (11)) and $M_{\text{line, vir}}$ (Eq. (12)) values of the substructures. The densities and the masses are derived from the continuum observations. The values are computed over the areas shown in Fig. 13, assuming a distance of 620 pc. For consistency with the clump parameters in Table 3, the average velocity dispersion is computed from the ^{13}CO data.

3.5. Molecular spectra of compact sources

For each clump, we computed average spectra for each detected molecular line. The spectra were computed by taking an average over $50'' \times 50''$, centred at the peak of the column density and averaging over a velocity interval of 0.125 km s^{-1} . The spectra are shown in Figs. 14 and 15. The red and black dashed lines in Fig. 14 show the velocity range used for mass accretion calculation in Sect. 5.

The peak of the emission is seen in the range $[3, 5] \text{ km s}^{-1}$ except for the clump B where the peak is at $\sim 2.5 \text{ km s}^{-1}$. For each clump, the ^{12}CO line has broader features compared to the other molecular lines, often showing two or even three peaks and strong asymmetry. The complexity of the ^{12}CO line may arise from several velocity components along the line-of-sight as can be seen in Fig. 5, and/or from self-absorption. The blue-shifted part of ^{13}CO and C^{18}O spectra follow roughly the same

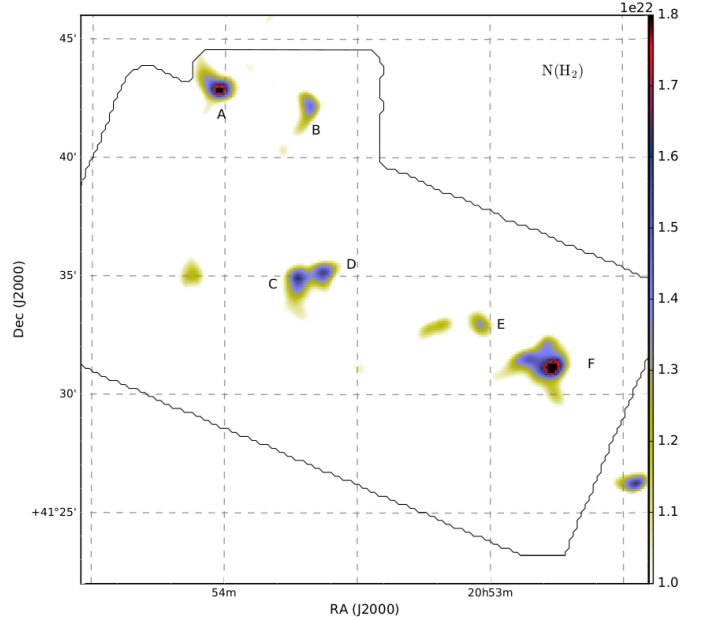


Fig. 12. Masked H_2 column-density map, showing the locations and column density of the clumps (labelled A to F) for which we have computed a mass estimate. The black contour shows the area of the Nobeyama observations.

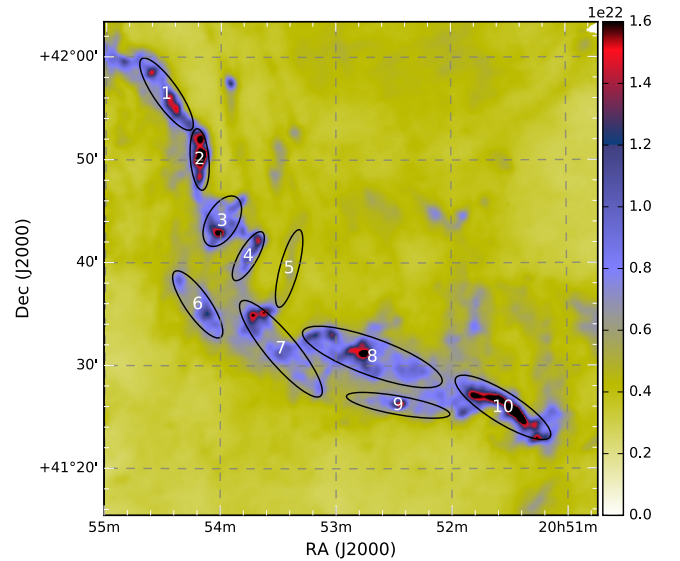


Fig. 13. Locations and surface areas of the identified substructures (labelled 1 to 10). The underlying colour map is the column density of H_2 derived from the dust observations.

shape as ^{12}CO , although for most clumps the peak of ^{13}CO and C^{18}O emission is between the two peaks of the ^{12}CO spectra. In Fig. 14, the HCO^+ and CS lines show blue asymmetric profiles (negative velocity shift by $\sim 0.5 \text{ km s}^{-1}$ compared to the CO lines) for clumps E and F, a possible indication of infall. For clump A the HCO^+ and CS lines have a positive velocity shift. For the remaining clumps, the HCO^+ and CS have velocities consistent with the rarer CO isotopomers.

The spectra taken from the clump B show a clear double-peaked profile in ^{13}CO and possibly in C^{18}O . The strongest HCO^+ , HCN, and CS emission is also seen at clump B, an indication of a different chemical composition compared to the other clumps. The HCN emission is also detected at clumps A, C, E,

Table 3. Average hydrogen column density, average density, velocity dispersion, estimated clump mass, and virial mass for the clumps shown in Fig. 12.

Clump	$N(\text{H}_2)^1$ (10^{22} cm^{-2})	$n(\text{H}_2)^1$ (10^4 cm^{-3})	$\sigma_{^{13}\text{CO}}$ (km s^{-1})	$\sigma_{\text{C}^{18}\text{O}}$ (km s^{-1})	M_c^1 (M_\odot)	$M_{\text{virial}}(^{13}\text{CO})$ (M_\odot)	$M_{\text{virial}}(\text{C}^{18}\text{O})$ (M_\odot)
A	1.54 ± 0.04	1.93 ± 0.05	0.68 ± 0.02	0.53 ± 0.01	15.2 ± 0.41	18.5 ± 0.94	11.4 ± 0.94
B	1.31 ± 0.03	2.07 ± 0.06	0.61 ± 0.03	0.52 ± 0.02	8.1 ± 0.22	11.7 ± 0.62	8.6 ± 0.94
C	1.46 ± 0.05	2.11 ± 0.05	0.88 ± 0.03	0.61 ± 0.03	10.6 ± 0.38	27.0 ± 0.84	15.6 ± 0.94
D	1.44 ± 0.04	2.26 ± 0.07	0.58 ± 0.02	0.41 ± 0.01	8.8 ± 0.36	10.8 ± 0.71	5.4 ± 0.94
E	1.25 ± 0.02	2.68 ± 0.08	0.61 ± 0.01	0.42 ± 0.02	4.2 ± 0.17	8.7 ± 0.53	4.3 ± 0.94
F	1.65 ± 0.04	1.93 ± 0.04	0.62 ± 0.02	0.51 ± 0.02	18.5 ± 0.52	16.4 ± 0.83	11.3 ± 0.94

Notes. A distance of 620 pc is assumed. ⁽¹⁾ Estimated from dust emission.

Table 4. Properties of the selected substructures identified in Fig. 13, assuming a distance of 620 pc.

Substructure	$N(\text{H}_2)^1$ (10^{21} cm^{-2})	$n(\text{H}_2)^1$ (10^3 cm^{-3})	$\sigma_{^{13}\text{CO}}$ (km s^{-1})	M_s^1 (M_\odot)	M_{line}^1 (M_\odot/pc)	$M_{\text{line,vir}}$ (M_\odot/pc)
1	7.99 ± 0.27	6.40 ± 0.07	0.69 ± 0.02	76.2 ± 3.5	94.2 ± 2.9	86.2 ± 3.4
2	11.2 ± 0.58	12.3 ± 0.11	0.71 ± 0.03	57.4 ± 2.3	96.9 ± 4.0	221.1 ± 4.4
3	7.91 ± 0.28	7.06 ± 0.08	0.69 ± 0.02	60.7 ± 4.8	83.6 ± 6.9	193.7 ± 4.7
4	7.31 ± 0.49	8.42 ± 0.10	0.73 ± 0.02	33.6 ± 3.3	59.8 ± 3.9	221.9 ± 3.6
5	3.66 ± 0.14	3.51 ± 0.07	0.43 ± 0.01	24.3 ± 4.4	36.0 ± 2.9	84.7 ± 3.9
6	6.50 ± 0.36	5.40 ± 0.06	0.82 ± 0.02	57.6 ± 4.1	73.8 ± 3.6	283.3 ± 4.9
7	6.94 ± 0.27	4.16 ± 0.08	0.79 ± 0.03	118.2 ± 4.6	109.3 ± 4.3	248.7 ± 7.2
8	7.50 ± 0.32	3.74 ± 0.04	0.58 ± 0.02	183.9 ± 7.0	141.6 ± 5.4	178.5 ± 8.5
9	6.09 ± 0.62	5.11 ± 0.03	0.56 ± 0.01	52.9 ± 4.2	68.5 ± 2.9	142.3 ± 2.8
10	9.65 ± 0.29	6.09 ± 0.05	0.58 ± 0.03	148.0 ± 9.3	144.2 ± 5.3	125.9 ± 7.2

Notes. ⁽¹⁾ Estimated from dust emission.

and F (see Fig. 15). We performed a hyperfine fit to the spectra, but only obtained meaningful results for the clump B due to a low S/N (see Fig. C.1). The inferred excitation temperature of HCN and the optical depth of the emission line are 6.0 K and 2.0, respectively.

The ^{13}CO spectra taken at clumps C and D show a clear tail towards lower velocities, indicating a possible outflow or an extended foreground (or background) layer towards both clumps. Clumps A and D also show the strongest SO emission (see Fig. 15).

The ^{12}CO velocity maps from both Osaka and Nobeyama observations show a distinct velocity component in the range [6.5, 9.5] km s^{-1} , elongating orthogonal to the main filament (see Fig. A.1). To study the kinematics of the component, we have extracted spectra from positions L1 and L2 (see Fig. 6), and displayed them in Fig. D.1. The spectra show a complex velocity structure, but the strongest emission is still seen in the range of 3–4 km s^{-1} , with a secondary maximum between 6 and 8 km s^{-1} .

3.6. Fragmentation and nonthermal motions

The *Herschel* observations show that the cold filament is fragmented to several substructures. In order to estimate a fragmentation scale length along the cold filament, we have computed an average hydrogen column density following the ridge of the cold filament. Shown in the right panel of Fig. 16 is the density distribution along the white line in the left panel starting from the upper left corner. There seems to be two fragmentation

scales, a small-scale fragmentation between individual clumps of ~ 0.25 pc, and a large-scale fragmentation at ~ 0.8 pc. The small-scale fragmentation length is close to the Jeans' length (Chandrasekhar & Fermi 1953)

$$\lambda_J = \sqrt{\frac{\pi c_s^2}{G\rho}}, \quad (13)$$

where c_s is the sound speed, G is the gravitational constant and ρ is the density. Inserting $c_s = 0.21 \text{ km s}^{-1}$ (assuming a temperature of 15 K) and $\rho = 2.16 \times 10^3 \text{ cm}^{-3}$ (average over the values in Table 4) results in $\lambda_J = 0.3$ pc.

On the other hand, at larger scales, the fragmentation length follows more closely the fragmentation model of infinitely long cylinder (Inutsuka & Miyama 1992; Jackson et al. 2010), where the separation length λ_{max} depends on the filament scale height $H = \sqrt{\pi c_s^2 / G\rho}$ and the radius R of the filament. If $R \gg H$, the fragmentation length is $\lambda_{\text{max}} = 22H$. Using the values above for c_s and ρ , the separation length is $\lambda_{\text{max}} = 1.05$ pc.

The nonthermal motions of the gas can be associated to turbulence or core formation. To study the nonthermal contribution to the line width, the thermal component has to be subtracted. We assume that the two components are independent and define the nonthermal velocity dispersion as

$$\sigma_{\text{nt}} = \sqrt{\frac{\Delta V^2}{8 \ln 2} - \frac{k_b T_{\text{kin}}}{m}}. \quad (14)$$

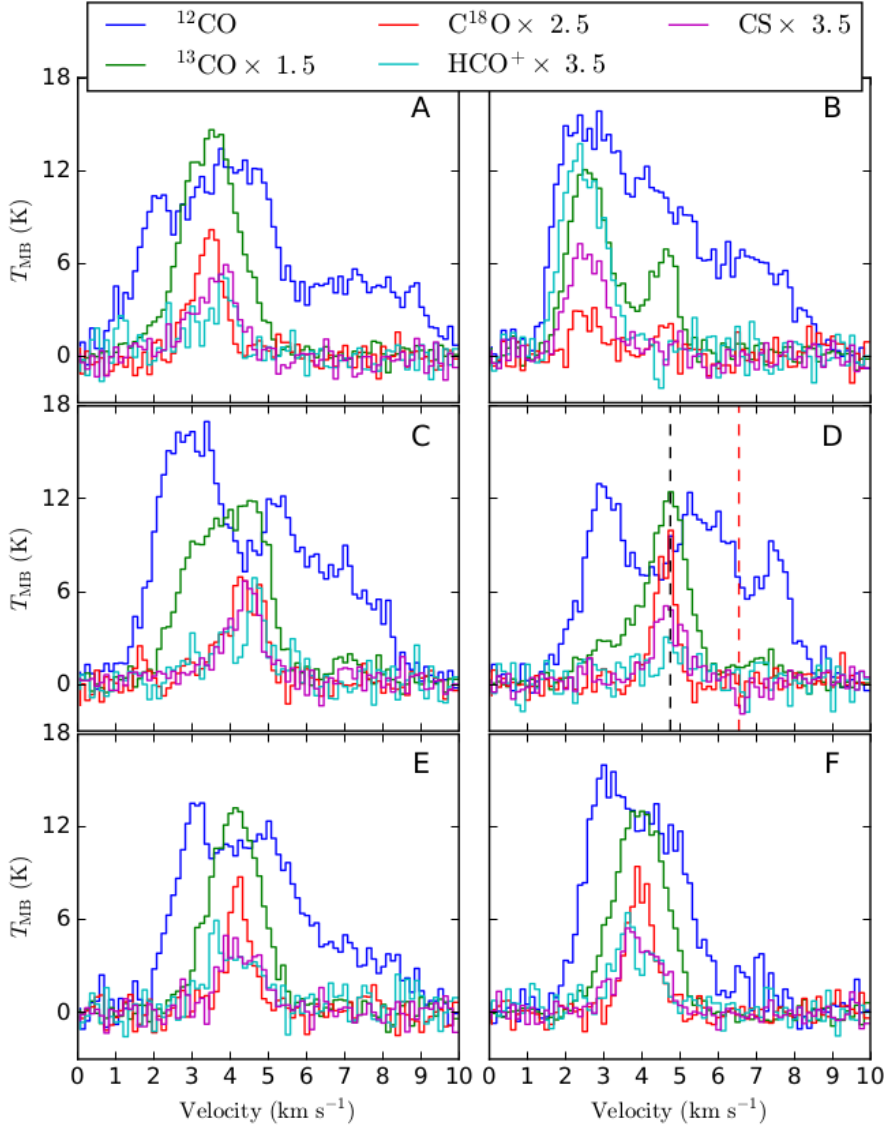


Fig. 14. Spectra detected from clumps A–F in Fig. 6. The spectra are scaled by the factors shown at the top of the figure. The red and black broken lines in the panel for clump D indicate the velocity range used to calculate the mass accretion.

The nonthermal velocity dispersion can be compared with the sound speed of the gas c_s , for which we have used a value of 0.21 km s^{-1} , corresponding to a gas temperature of 15 K.

Shown in Fig. 17 are the histograms of the velocity dispersion of SO, ^{13}CO , and C^{18}O . For the histograms, only the pixels with $S/N > 4.0$, computed from the total integrated intensity, were included. Only a few points have subsonic velocity dispersion, and an overwhelming majority have supersonic motions. The subsonic regions are located close to or between the outer regions of clumps B, D, and E (Fig. 6).

The linewidth-size relationship, one of the Larson’s relations (Larson 1981), defines the correlation between the velocity dispersion of the gas as a function of the scale. Shown in the left panel of Fig. 18 is the relation derived from the Osaka ^{13}CO data. The data points in the panel were computed by dividing the cold filament to length bins, starting from the north-western part of the cold filament and moving south-east along the filament. For the first point, the length of each bin is 0.25 pc, and for each subsequent data point the length of the bins was increased by 0.25 pc, up to 9.0 pc. For each data point in the figure, the velocity dispersion is computed as an average over the bins.

As seen in the left panel of Fig. 18, the velocity dispersion as a function of the size scale is relatively uniform. However, at the scale of one parsec, the scale of individual cloud fragments, the dispersion shows large variations along the filament (see right panel of Figs. 18 and 16).

4. Discussion

4.1. Colliding filaments

The line observations (Figs. B.1 and B.2) indicate that the velocity field around the central part of the filament, between clumps C and F (right panel of Fig. 6), is not uniform. The large-scale images from *Planck*, WISE, and *Herschel* PACS instrument (see Figs. 2 and 3) show a warm structure, which is seen faintly in the *Herschel* image (Fig. 1) as a red filamentary structure running from south-east to north-west across the cold filament. The same warm filament can be seen in both the Osaka and the Nobeyama ^{12}CO observations in the range of $\sim[7, 9] \text{ km s}^{-1}$ (Figs. A.1 and B.1, respectively). Furthermore, the Osaka

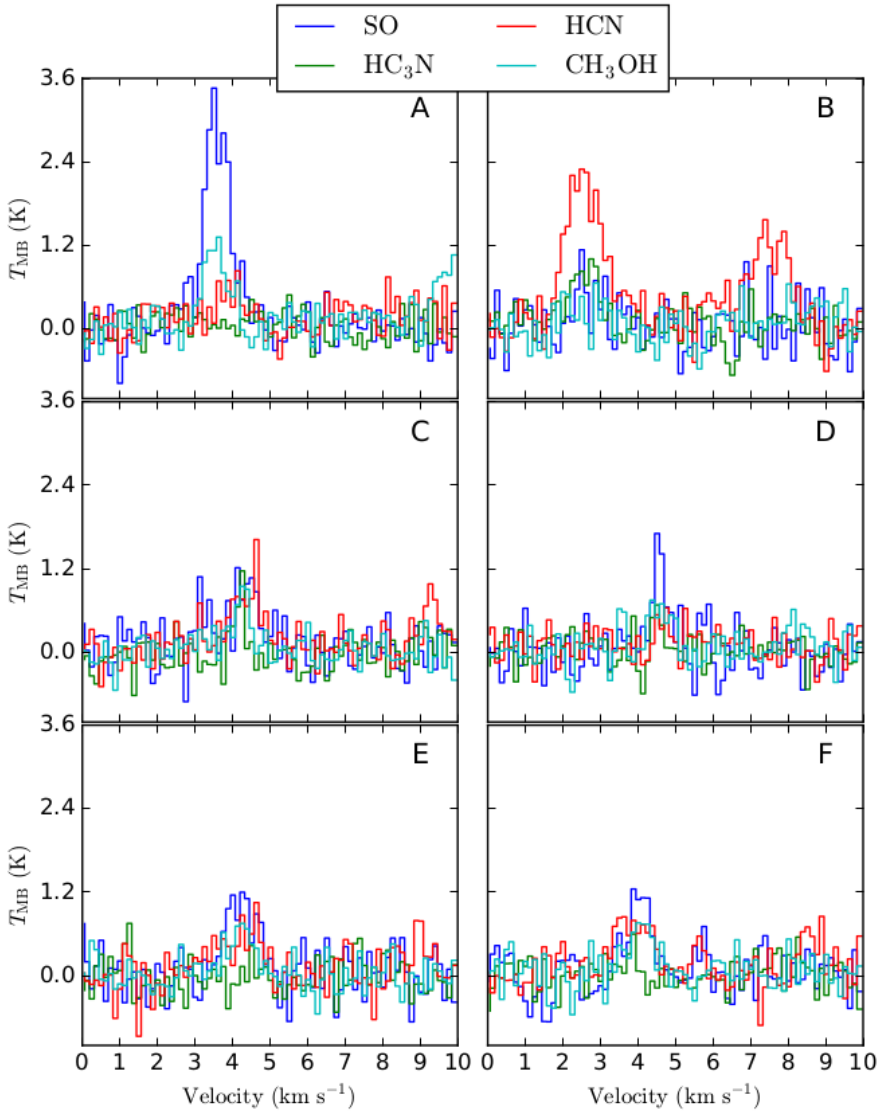


Fig. 15. Spectra of the weaker lines for the clumps shown in Fig. 6. We note that the secondary peak at ~ 7 km s $^{-1}$ in the HCN spectrum of the clump B is one of the hyperfine components.

^{12}CO observations in the range of $\sim[-4, -1]$ km s $^{-1}$ show a separate velocity component at the western end of the mapped region, which is not associated to either the cold or the warm filament.

The complex velocity field as well as the temperature difference of the cold and warm filaments suggest a possibility that the two filaments are distinct structures without physical connection lying along the same line-of-sight. The distance estimates are 620 pc for the cold filament and, depending on the mask used, 300 pc or 670 pc for the warm filament. Thus, if the warm filament is truly ~ 300 pc closer than the cold filament, the temperature difference can be explained by differences in the local radiation fields. However, if the warm filament is at a distance similar to the cold filament, the temperature difference should represent the difference in density, that is, the cold filament is denser than the warm filament.

Looking at the *Herschel* optical depth map (right panel of Fig. 6), there are two cavities, one between clumps B and D, and the other between clumps D and E. Interestingly, the cavities seem to be aligned roughly along the warm filament (see Fig. 1). If the distance to the warm filament is 670 pc, it would seem that

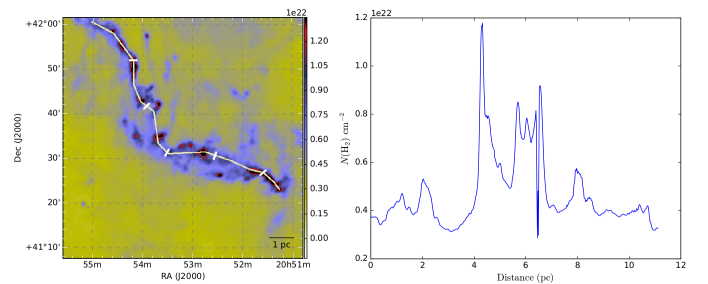


Fig. 16. *Right panel:* average hydrogen column density following the white line in the *left panel*. The *x*-axis of the *right panel* shows the distance along the white line, starting from the upper-left corner and the white ticks on the line show steps of 2 pc. The column density values are computed as an average over 3×3 pixels. The colour map of the *left panel* corresponds to the hydrogen column density derived from the dust observations.

clumps D and C (possibly also clump B) are close to areas where a collision between two structures is possible.

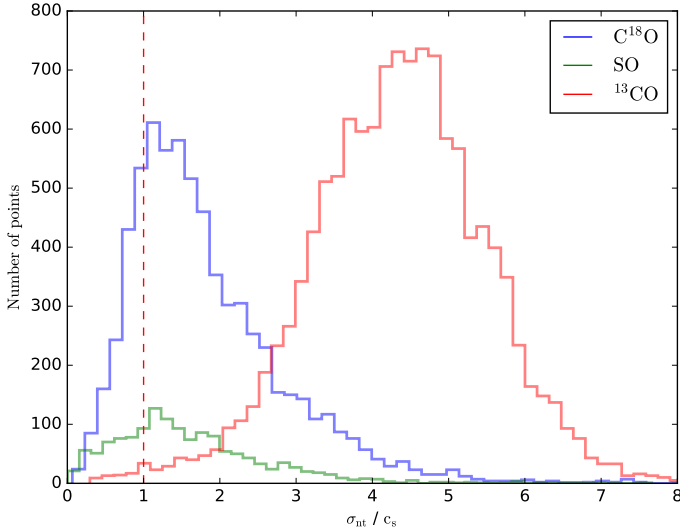


Fig. 17. Ratio of nonthermal velocity dispersion (σ_{nt}) over sound speed (c_s) for SO, $^{13}\text{CO}_{\text{NO}}$, and $\text{C}^{18}\text{O}_{\text{NO}}$. The red line represents the limit $\sigma_{\text{nt}}/c_s = 1.0$.

SO can be used as a tracer molecule for shocks and outflows (Pineau des Forets et al. 1993; Jørgensen et al. 2004), thus the SO emission from the two clumps, D and B, can indicate that the clumps have formed as a result of colliding structures. In addition, both clumps show CH_3OH emission in the same velocity interval as SO emission (Figs. B.4 and B.8), which could be an indication of outflows or shocks (Jørgensen et al. 2004). The SO contours in Fig. B.8 seem to trace the inner edges of the two cavities, indicating a possible shock. However, for both SO and CH_3OH , the S/N is low, and more sensitive observations are needed.

A collision between the cold filament and the warm filament is possible if the two components are at a similar distance. In such an event, most of the clumps identified from the central part of the main filament are close to the shock front, and thus could be formed as a result of the collision.

If the warm filament is located at ~ 300 pc, the collisional model is untenable. Nevertheless, the *Herschel* observations show that the central region of the cold filament is significantly broader compared to the northern and southern ends of the filament, indicating that the filament is interacting with the surrounding medium. One possible explanation could be that the entire filament is moving north-west with a higher relative velocity relative to the surrounding medium, or is interacting with radiation from nearby bright stars.

4.2. Striations

Several papers, for example Goldsmith et al. (2008) and Palmeirim et al. (2013), have studied the relationship between striations and the parent filament. Some recent studies have shown a close connection between the striations and the large-scale magnetic field (e.g. Malinen et al. 2016; Cox et al. 2016). The *Herschel* observations of G82.65-2.00 clearly show a host of faint structures around the main filament. For the purposes of this work we identify all faint and well defined elongated structures as striations.

The striations raise a question about the dust properties, as the striations are faint but clearly appear in blue in Fig. 1,

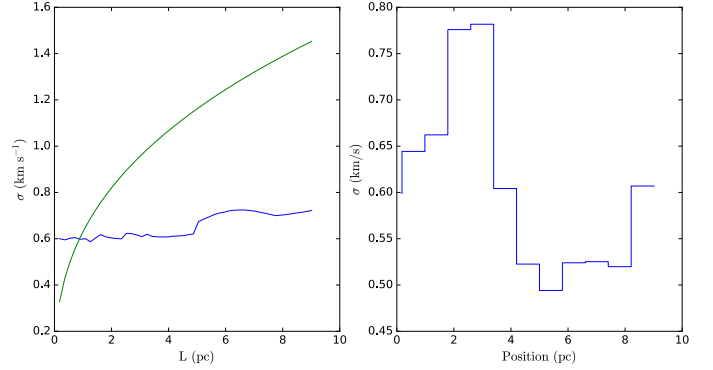


Fig. 18. *Left panel:* correlation of the velocity dispersion as a function of the scale, the blue line, derived using the Osaka ^{13}CO data. The green curve is the Larson's relation $\sigma = 0.63 L^{0.38}$, where L is the scale length; and *right panel:* velocity dispersion as a function of position in the filament, starting from north-west and moving south-east along the filament at one parsec scale.

and thus are clearly separated from the warm background. The average column density of the striation S1, listed in Table 4, is lower by a factor of ~ 2 compared to the rest of the filament. However, the dust seems to be almost as cold in the striations as in the cold filament and clearly colder than in the general background that mainly corresponds to the warm filament visible in, for example, the $12\ \mu\text{m}$ image (see Fig. 2). Thus, either the striations are not connected to the cold filament and are in a different radiation field, or the properties of the dust grains (mainly the size of the grains) have changed.

However, the ^{13}CO and C^{18}O observations from Nobeyama connect one of the striations, S1, to the cold filament. Indeed, the ^{13}CO and C^{18}O (Figs. B.2 and B.3) contours around $\sim 5\ \text{km s}^{-1}$ trace the striation S1 and the two clumps, C and D, in the central part of the filament. Furthermore, the PV cut along the striation S1 shows a continuous velocity structure. Thus, it is likely that the striation is connected to the main filament and might be an indication of matter accretion on the filament. On the other hand, because of the active star formation or because of some external forces, and if the cold colour temperature of the striations is a sign of larger dust grains, this might also be material stripped away from the filament.

We can derive a rough estimate for the potential mass accretion from the striation S1 onto the clump D, by assuming a simple relation for a core accreting mass along a cylindrical structure

$$\dot{M} = \frac{M_s \hat{V}}{L_s}, \quad (15)$$

where M_s is the mass of the striation, L_s is the length of the striation, and \hat{V} is the velocity. As the value of \hat{V} , we use the difference between the central velocity of the ^{13}CO line (black dashed line in Fig. 14) and the extent of the red lobe of ^{12}CO line (red dashed line in Fig. 14). Inserting $L_s = 2.17$ pc, $\hat{V} = 1.7\ \text{km s}^{-1}$, and $M = 36.0 M_\odot$, we obtain an estimate for the mass accretion rate of $\dot{M} = 2.2 \times 10^{-6} M_\odot/\text{yr}$.

In a recent study López-Sepulcre et al. (2010) analysed a selection of 48 high-mass molecular clumps, 10 of which were inferred as infall candidates. The mass accretion rates for the 10 sources were derived from the outflow mass-loss rates (for a review of the method see Beuther et al. 2002; and López-Sepulcre et al. 2009) resulting in accretion rates of $\sim 10^{-5} M_\odot/\text{yr}$. Similarly, Palmeirim et al. (2013) studied the

Taurus B211/3 filament identifying several striations that seem to be connected to the main filament. [Palmeirim et al. \(2013\)](#) explored several methods to derive the infall velocity, deriving accretion rates between 2.7 and $5.0 \times 10^{-5} M_{\odot}/\text{yr}$ from a $220 M_{\odot}$ striation onto the main filament and adopting a velocity of $\sim 1 \text{ km s}^{-1}$.

Comparing the above numbers, our derived value is lower by an order of magnitude. However, the environments and accretion models under consideration differ significantly and because our clumps are smaller, a smaller accretion rate seems reasonable.

4.3. Kinematics

The M_{line} values for almost all of the substructures are close to $\sim 100 M_{\odot}/\text{pc}$, which seems to hold also for larger scales. The estimated total mass of the filament is $\sim 1200 M_{\odot}$ and the length of the filament is $\sim 10 \text{ pc}$. Thus, the resulting M_{line} for the main filament is $\sim 120 M_{\odot}/\text{pc}$. Only three of the substructures (Nos. 1, 8, and 10 in Table 4) have M_{line} greater than or close to $M_{\text{line, vir}}$. The result indicates that most of the structures are not bound and will eventually disperse back to the ISM and only the three substructures would be able to form new stars.

However, the derived virial masses might be overestimated. The velocity dispersions used for computing the $M_{\text{line, vir}}$ are derived from the large-scale Osaka ^{13}CO observations. If the ^{13}CO line is saturated, the derived velocity dispersion will be higher because of the broader line profile. The pixel-to-pixel correlation between the τ_{250} and the $N(^{13}\text{CO})$, Fig. 9, indicates that the ^{13}CO line is indeed saturated. Thus, the velocity dispersions should be estimated from the C^{18}O line. However, because of the lower S/N of the C^{18}O data and because we do not detect C^{18}O in the large-scale observations, we use ^{13}CO . On the other hand, because of the high column density, the derived velocity dispersion might not give a proper picture of the inner motion of the clumps or substructures. Furthermore, the lower resolution will increase the uncertainty of the derived values. An uncertainty of 5% in the line width will change the derived $M_{\text{line, vir}}$ value by several M_{\odot} . The M_{line} and $M_{\text{line, vir}}$ values were computed with the assumption of isothermal cylinder. The assumption is probably not an accurate description for most of the substructures as they have already formed cores as can be seen in Fig. 6.

The $N(^{13}\text{CO})$ line was also used to estimate the virial masses of the cold clumps (Table 3) with the exception that we used the higher-resolution Nobeyama observations. For most clumps, the estimated virial masses are larger than the clump masses, however, for the clumps we computed a second estimate for the velocity dispersion using the C^{18}O line which is not showing signs of saturation. The velocity dispersions measured in C^{18}O are smaller, as summarised in Table 3, yielding smaller virial masses, indicating that most of the clumps have masses larger than or close to the virial masses. Thus it is possible that some of the substructures might be bound.

As discussed by [Arzoumanian et al. \(2013\)](#), a higher velocity dispersion can be related to mass accretion. The higher velocity dispersion values we derive for substructures 4 and 7, and possibly substructure 6, might thus be related to mass accretion, as the structures seem to be connected to striations. In such a case, the higher-velocity dispersion is not caused by larger-scale turbulence but rather by amplification of initial velocity fluctuations via conversion from gravitational energy to kinetic energy during mass accretion ([Arzoumanian et al. 2013](#)). However, due to limited coverage of our higher-resolution observations, the connection between the substructures and the striations cannot be confirmed.

The recent findings of [Hacar & Tafalla \(2011\)](#) and [Hacar et al. \(2016\)](#), in L1517 and in Musca, show two filaments that are clearly subsonic, and, at least in Musca, do not follow the Larson's velocity dispersion-size relationship. We see, in Fig. 18, that most of the gas, at least in the central regions, of G82.65-2.00 is clearly supersonic, with $\sigma_{\text{nt}}/\sigma_s > 1.0$. However, the derived Larson's relation is almost flat, indicating that the overall velocity dispersion along the filament, even at larger scales, is relatively constant. However, at the 1 pc scale the substructure of the larger filament is clearly seen. Thus, as the filament is likely dispersing, the dispersion is occurring slowly.

5. Conclusions

The *Herschel* field G82.65-2.00 is a filamentary field at a late stage of evolution, a “debris filament”. The central region of the cold, main filament has a clearly ragged morphology, and has formed several cold cores. Furthermore, the presence of several YSOs and the relatively high-velocity differences along the cold filament can indicate that the filament is dispersing to the ISM. Our main conclusions are

1. Although our line observations agree well with the results of [Dame et al. \(2001\)](#), who inferred a distance of 1 kpc, different methods result in inconsistent distance estimates. We have derived a new distance estimate of $620 \text{ pc} \pm 40 \text{ pc}$ using an infrared extinction method.
2. The ^{12}CO maps both from Nobeyama and Osaka combined with *Herschel* and WISE observations show a separate, diffuse warm filament extending over the cold filament. Using the extinction method, we have derived a distance estimate of 300 or 670 pc for the warm filament, depending on the mask used, possibly placing it at a similar distance to the cold filament. Comparing the *Herschel* images and SO velocity contours, and assuming the two filaments are at a similar distance, it is possible that the two structures are colliding. Thus, most of the cold clumps in the central region of the cold filament could have been formed as a result of the collision.
3. The *Herschel* observations and the line observations from Nobeyama indicate that the average molecular hydrogen column density of the filament is in the range of $1.0\text{--}2.5 \times 10^{22} \text{ cm}^{-2}$. We have chosen six clumps for further analysis from the central region of the cold filament. The masses of the clumps vary in the range $10\text{--}20 M_{\odot}$ and the total mass of the clumps is $\sim 70 M_{\odot}$. Virial masses estimated from the velocity dispersion of the C^{18}O emission line are close to or slightly smaller than the clump masses. The chemical composition of the clumps is uniform, except for clump B for which the molecular line spectra shows distinct variations compared to the other clumps. We have visually identified substructures of the cold filament. The substructures have a wide range of masses from $\sim 20 M_{\odot}$ up to $\sim 180 M_{\odot}$, and in total, the mass of the cold filament exceeds $1200 M_{\odot}$. Only three of the substructures have masses close to or greater than the virial mass. However, the estimate is based on the ^{13}CO line, which is likely saturated in the densest regions. As in the case with the cold clumps, the derived virial masses would likely be smaller if the velocity dispersion could be estimated from the C^{18}O line.
4. Using the Nobeyama 45 m telescope, we have detected several kinematic components from the field. The strongest line

emission is seen in the velocity range [3, 5] km s⁻¹ with other notable emission around -4 and 8 km s⁻¹. The spectral lines show a complex velocity structure around the filament, and the velocity profiles of ¹²CO and ¹³CO along the striation S1 seem to indicate possible mass accretion from the striation on the filament. With a simple model we estimate a tentative mass accretion rate of $\dot{M} = 2.23 \times 10^{-6} M_{\odot}/\text{yr}$.

In order to study the fragmentation and possible dispersion of the filament in greater detail, more large-scale molecular line observations would be needed. A high-resolution map, for example in ammonia, would be useful to determine not only the substructure kinematics but also the state of the cold clumps. Furthermore, in connection with a more detailed study of the filamentary structures, the *Planck* polarisation data could be used to investigate the relation between magnetic fields and the identified substructures.

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Appendix A: Osaka velocity maps

Figures A.1 and A.2 show the velocity channel maps of ^{12}CO and ^{13}CO derived from the Osaka observations. For both velocity

maps, the underlying colour map is the $250\ \mu\text{m}$ optical depth. We do not display the C^{18}O channel map because of its low S/N.

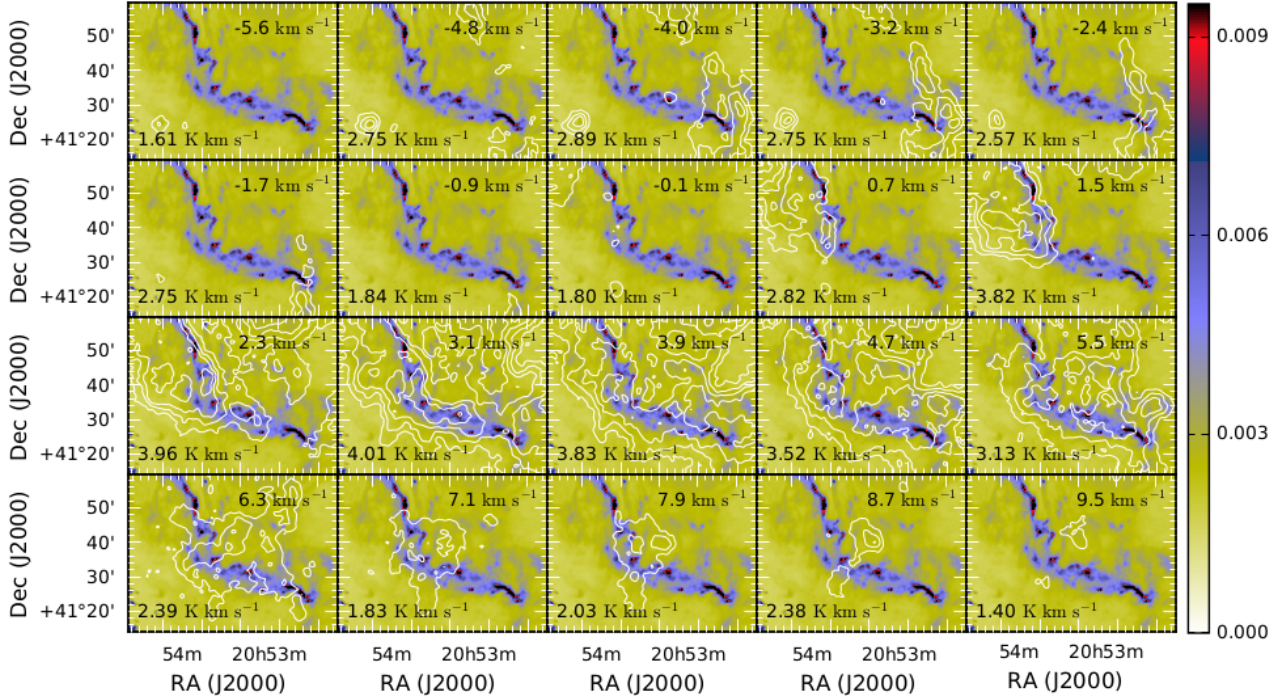


Fig. A.1. Osaka ^{12}CO line observations integrated over $0.79\ \text{km s}^{-1}$ wide velocity intervals. The average noise per $0.79\ \text{km s}^{-1}$ wide velocity interval is $\sigma = 0.40\ \text{K km s}^{-1}$. The lowest contour corresponds to 3σ , and the contours increase with 3σ steps. The numbers noted in the upper-right corner of each panel correspond to the centre of the velocity bin, and the numbers in the lower-left corner of each panel show the maximum intensity of the panel. The underlying colour map is the τ_{250} map.

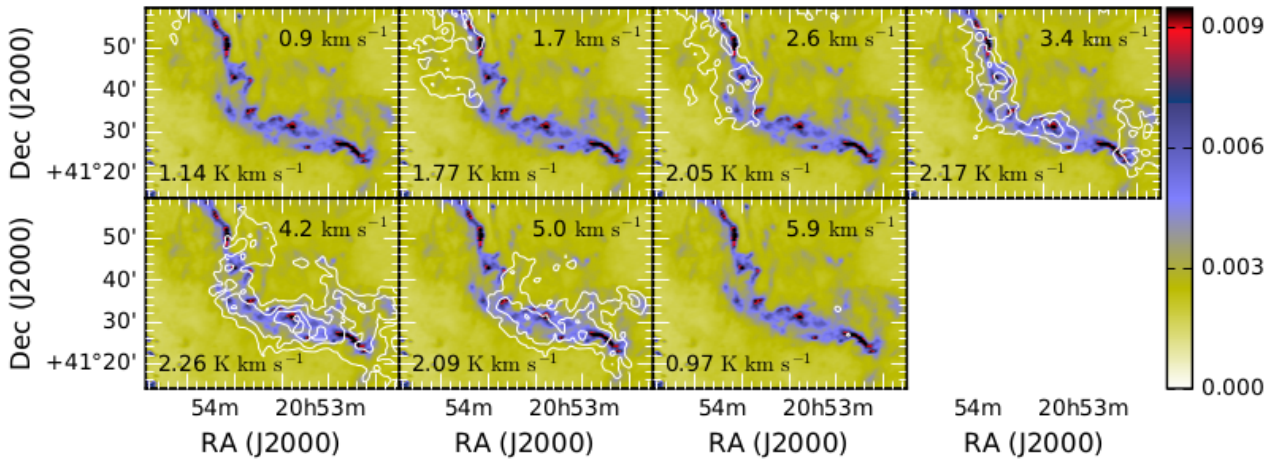


Fig. A.2. Same as Fig. A.1, but for ^{13}CO and integrated over $0.83\ \text{km s}^{-1}$ wide velocity intervals. The average noise per $0.83\ \text{km s}^{-1}$ wide velocity interval is $\sigma = 0.26\ \text{K km s}^{-1}$. The lowest contour corresponds to 3σ , and the contours increase with 2σ steps.

Appendix B: Nobeyama velocity maps

Figures B.1 to B.8 show the velocity channel maps for all of the molecular emission lines detected from the Nobeyama

observations. In all of the figures, the underlying colour map is the $250\ \mu\text{m}$ optical depth.

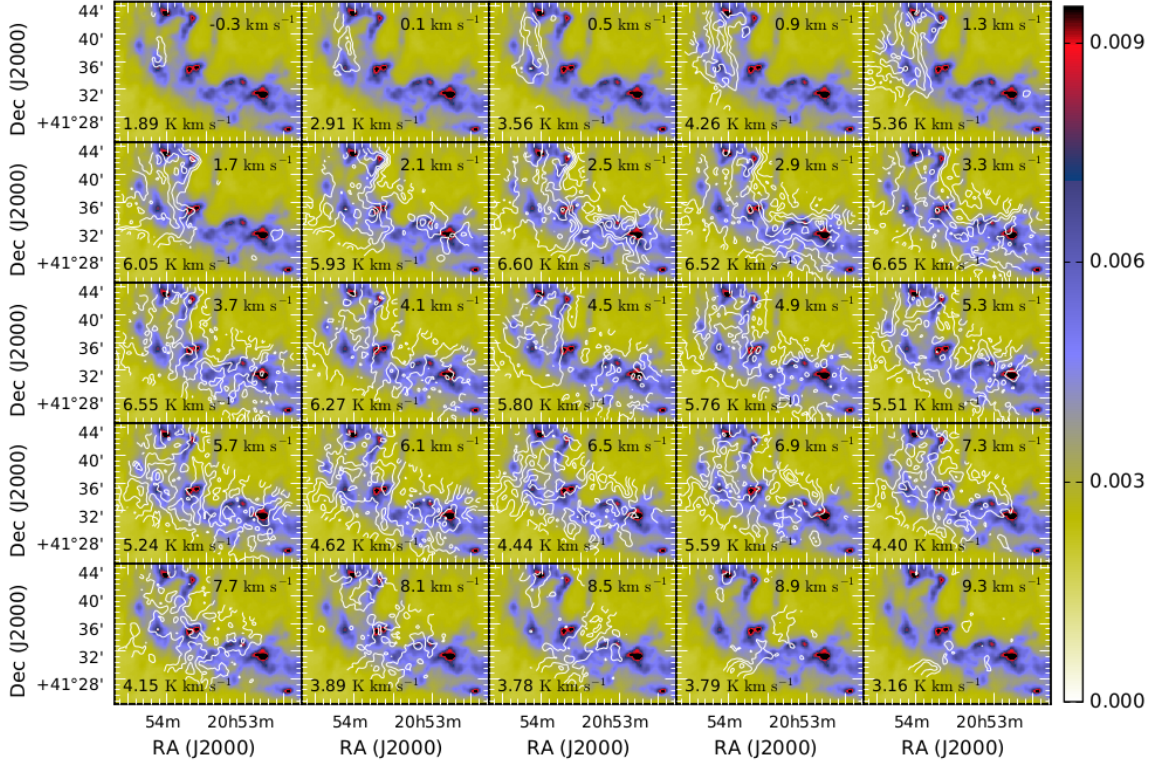


Fig. B.1. Integrated ^{12}CO line intensity W in $0.4\ \text{km s}^{-1}$ wide velocity intervals. The average noise per $0.4\ \text{km s}^{-1}$ wide velocity interval is $\sigma = 0.12\ \text{K km s}^{-1}$. The lowest contour corresponds to 6σ , and the contours increase with 6σ steps. The underlying colour map is the τ_{250} map. The bin centre velocity and the maximum value of W are given in the panels.

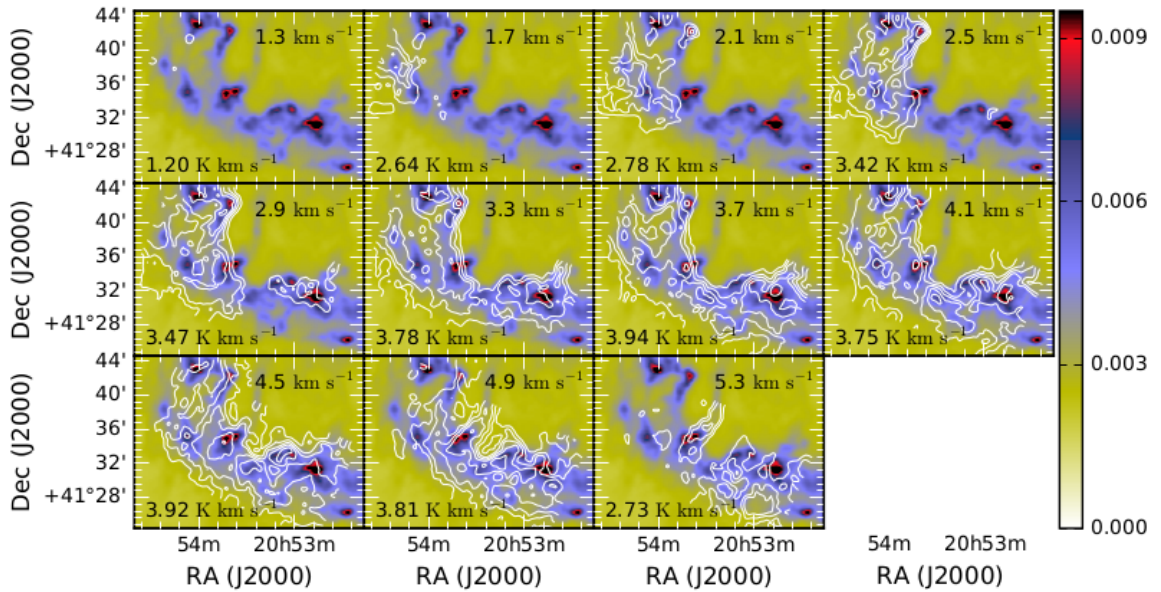


Fig. B.2. Same as Fig. B.1, but for ^{13}CO . The average noise per $0.4\ \text{km s}^{-1}$ wide velocity interval is $\sigma = 0.15\ \text{K km s}^{-1}$. The lowest contour corresponds to 6σ , and the contours increase with 4σ steps.

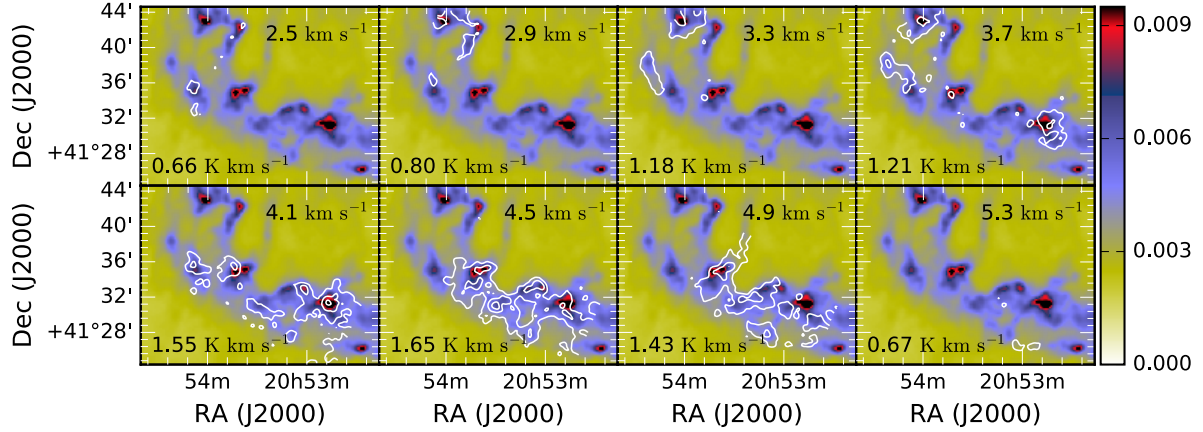


Fig. B.3. Same as Fig. B.1, but for C^{18}O . The average noise per 0.4 km s^{-1} wide velocity interval is $\sigma = 0.15 \text{ K km s}^{-1}$. The lowest contour corresponds to 3σ , and the contours increase with 3σ steps.

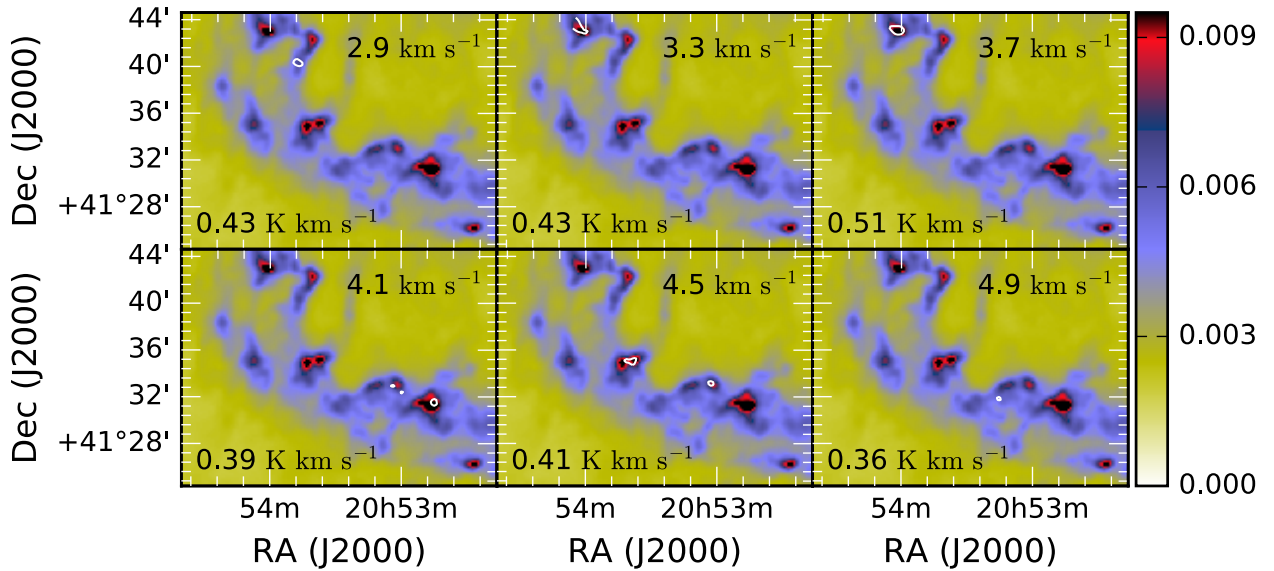


Fig. B.4. Same as Fig. B.1, but for CH_3OH . The average noise per 0.4 km s^{-1} wide velocity interval is $\sigma = 0.13 \text{ K km s}^{-1}$. The contour corresponds to 3σ .

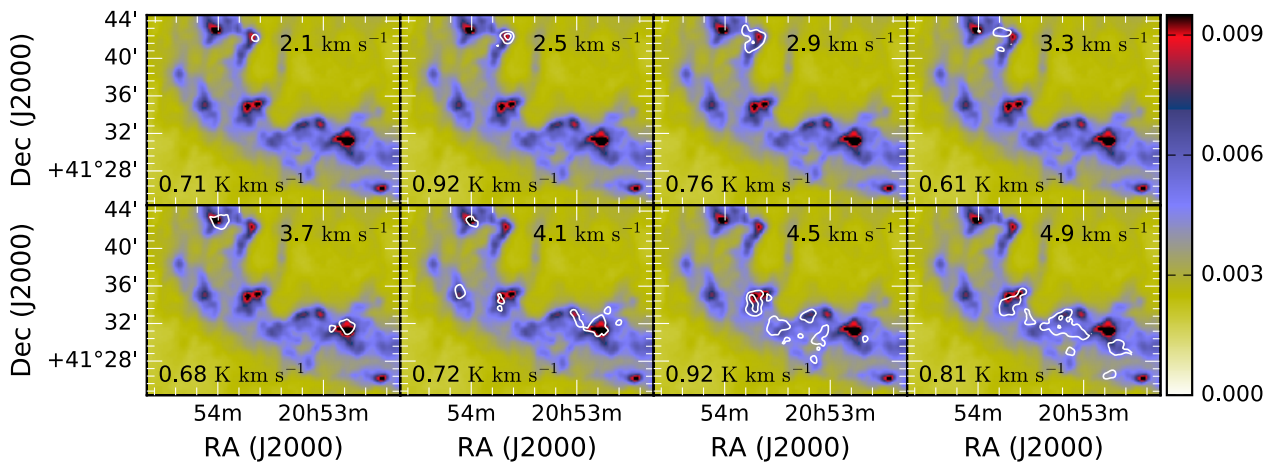


Fig. B.5. Same as Fig. B.1, but for CS . The average noise per 0.4 km s^{-1} wide velocity interval is $\sigma = 0.14 \text{ K km s}^{-1}$. The lowest contour corresponds to 3σ , and the contours increase with 2σ steps.

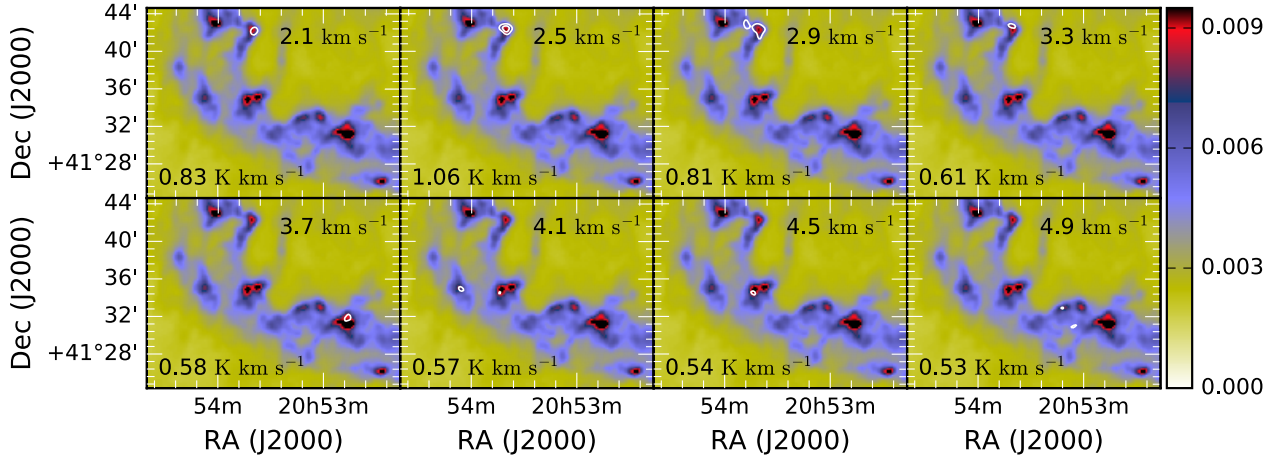


Fig. B.6. Same as Fig. B.1, but for HCN. The average noise per 0.4 km s^{-1} wide velocity interval is $\sigma = 0.16 \text{ K km s}^{-1}$. The lowest contour corresponds to 3σ , and the contours increase with 2σ steps.

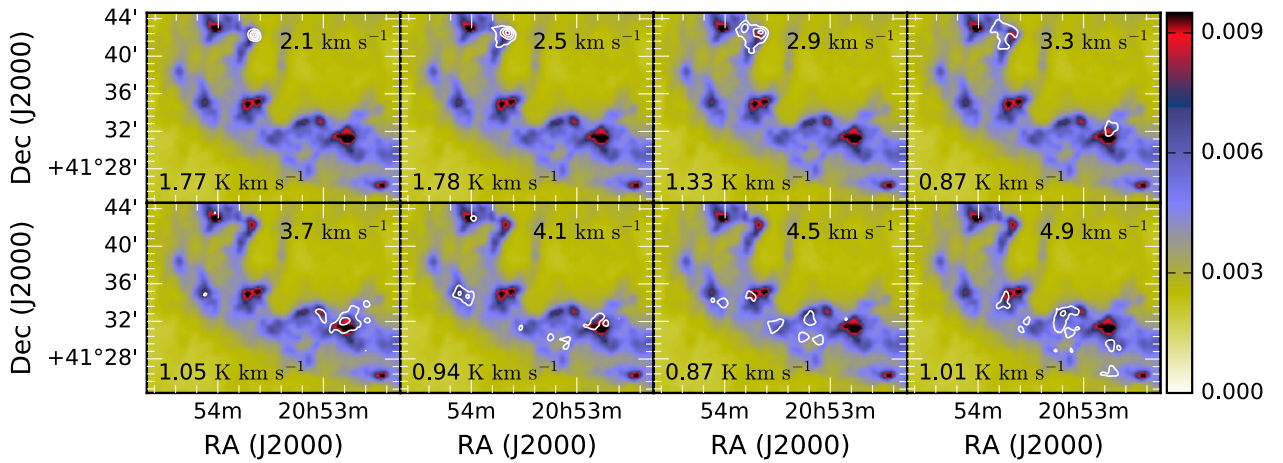


Fig. B.7. Same as Fig. B.1, but for HCO^+ . The average noise per 0.4 km s^{-1} wide velocity interval is $\sigma = 0.17 \text{ K km s}^{-1}$. The lowest contour corresponds to 3σ , and the contours increase with 2σ steps.

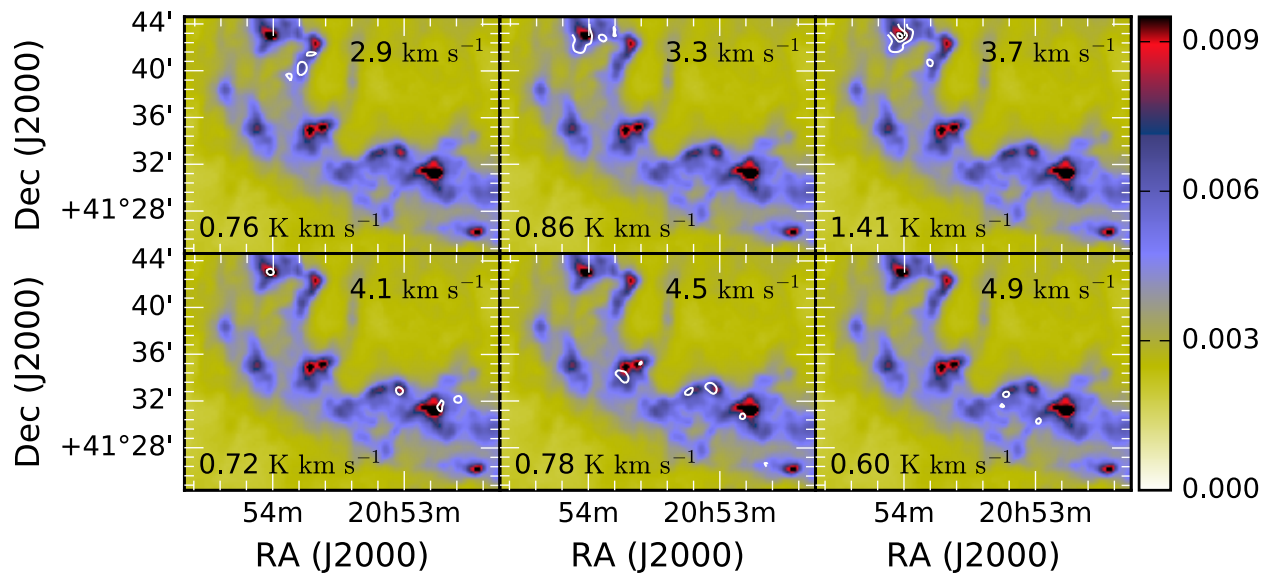


Fig. B.8. Same as Fig. B.1, but for SO. The average noise per 0.4 km s^{-1} wide velocity interval is $\sigma = 0.19 \text{ K km s}^{-1}$. The lowest contour corresponds to 3σ , and the contours increase with 2σ steps.

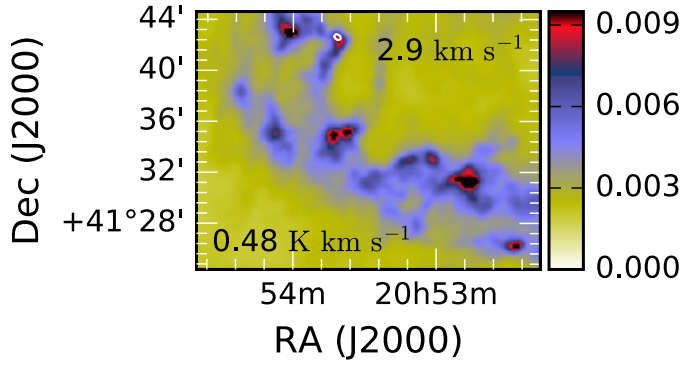


Fig. B.9. Same as Fig. B.1, but for HC_3N . The average noise per 0.4 km s^{-1} wide velocity interval is $\sigma = 0.15 \text{ K km s}^{-1}$. The contour corresponds to 3σ .

Appendix C: HCN hyperfine fitting

The HCN hyperfine structure can be used to derive additional information about the clump environments. We have used the `pyspeckit` package (Ginsburg & Mirocha 2011) for fitting the hyperfine structure of all six extracted spectra of HCN, from clumps A–F in Fig. 6. However, the S/N of the HCN lines was too low for all of the clumps to produce a reasonable fit, with the exception of clump B (shown in Fig. C.1).

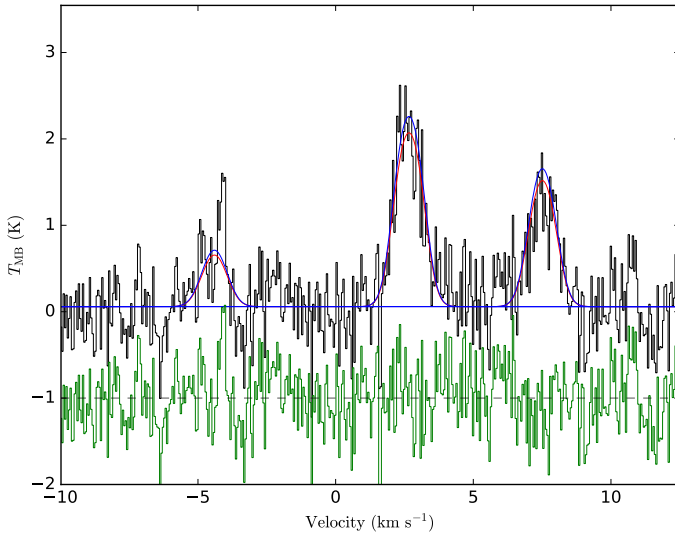


Fig. C.1. Hyperfine fit for the HCN spectra from clump B. The blue line is the result of hyperfine fitting and the red line is a simple Gaussian fit to the spectra. The green curve shows the residuals of the fitting. The residuals are vertically offset by -1 K .

Appendix D: Additional ^{12}CO spectra

Shown in Fig. D.1 are the additional Nobeyama ^{12}CO spectra extracted from positions L1 and L2 (see Fig. 6).

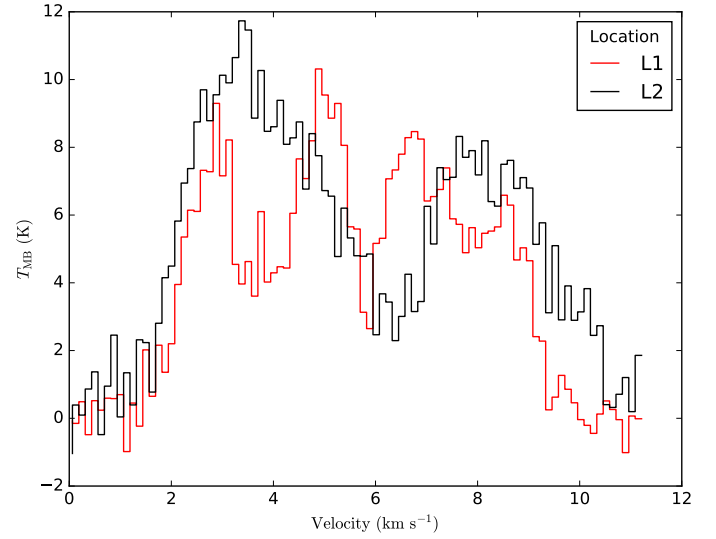


Fig. D.1. ^{12}CO ($J = 1-0$) spectra extracted from positions L1 and L2 in Fig. 6.