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Gaia Data Release 2

The celestial reference frame (*Gaia*-CRF2)

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ABSTRACT

Context. The second release of *Gaia* data (*Gaia* DR2) contains the astrometric parameters for more than half a million quasars. This set defines a kinematically non-rotating reference frame in the optical domain. A subset of these quasars have accurate VLBI positions that allow the axes of the reference frame to be aligned with the International Celestial Reference System (ICRF) radio frame.

Aims. We describe the astrometric and photometric properties of the quasars that were selected to represent the celestial reference frame of *Gaia* DR2 (*Gaia*-CRF2), and to compare the optical and radio positions for sources with accurate VLBI positions.

Methods. Descriptive statistics are used to characterise the overall properties of the quasar sample. Residual rotation and orientation errors and large-scale systematics are quantified by means of expansions in vector spherical harmonics. Positional differences are calculated relative to a prototype version of the forthcoming ICRF3.

Results. *Gaia*-CRF2 consists of the positions of a sample of 556 869 sources in *Gaia* DR2, obtained from a positional cross-match with the ICRF3-prototype and AllWISE AGN catalogues. The sample constitutes a clean, dense, and homogeneous set of extragalactic point sources in the magnitude range $G \approx 16$ to 21 mag with accurately known optical positions. The median positional uncertainty is 0.12 mas for $G < 18$ mag and 0.5 mas at $G = 20$ mag. Large-scale systematics are estimated to be in the range 20 to 30 μ as. The accuracy claims are supported by the parallaxes and proper motions of the quasars in *Gaia* DR2. The optical positions for a subset of 2820 sources in common with the ICRF3-prototype show very good overall agreement with the radio positions, but several tens of sources have significantly discrepant positions.

Conclusions. Based on less than 40% of the data expected from the nominal *Gaia* mission, *Gaia*-CRF2 is the first realisation of a non-rotating global optical reference frame that meets the ICRS prescriptions, meaning that it is built only on extragalactic sources. Its accuracy matches the current radio frame of the ICRF, but the density of sources in all parts of the sky is much higher, except along the Galactic equator.

Key words. astrometry – reference systems – catalogs

1. Introduction

One of the key science objectives of the European Space Agency's *Gaia* mission (Gaia Collaboration et al. 2016) is to build a rotation-free celestial reference frame in the visible wavelengths. This reference frame, which may be called

the *Gaia* Celestial Reference Frame (*Gaia*-CRF), should meet the specifications of the International Celestial Reference System (ICRS; Arias et al. 1995) in that its axes are fixed with respect to distant extragalactic objects, that is, to quasars. For continuity with existing reference frames and consistency

across the electromagnetic spectrum, the orientation of the axes should moreover coincide with the International Celestial Reference Frame (ICRF; Ma et al. 1998) that is established in the radio domain by means of VLBI observations of selected quasars.

The second release of data from *Gaia* (*Gaia* DR2; Gaia Collaboration 2018) provides complete astrometric data (positions, parallaxes, and proper motions) for more than 550 000 quasars. In the astrometric solution for *Gaia* DR2 (Lindgren 2018), subsets of these objects were used to avoid the rotation and align the axes with a prototype version of the forthcoming third realisation of the ICRF¹. The purpose of this paper is to characterise the resulting reference frame, *Gaia*-CRF2, by analysing the astrometric and photometric properties of quasars that are identified in *Gaia* DR2 from a positional cross-match with existing catalogues, including the ICRF3-prototype.

Gaia-CRF2 is the first optical realisation of a reference frame at sub-milliarcsecond (mas) precision, using a large number of extragalactic objects. With a mean density of more than ten quasars per square degree, it represents a more than 100-fold increase in the number of objects from the current realisation at radio wavelengths, the ICRF2 (Fey et al. 2015). The *Gaia*-CRF2 is bound to replace the HIPPARCOS Celestial Reference Frame (HCRF) as the most accurate representation of the ICRS at optical wavelengths until the next release of *Gaia* data. While the positions of the generally faint quasars constitute the primary realisation of *Gaia*-CRF2, the positions and proper motions of the ≈ 1.3 billion stars in *Gaia* DR2 are nominally in the same reference frame and thus provide a secondary realisation that covers the magnitude range $G \approx 6$ to 21 mag at similar precisions, which degrades with increasing distance from the reference epoch J2015.5. The properties of the stellar reference frame of *Gaia* DR2 are not discussed here.

This paper explains in Sect. 2 the selection of the *Gaia* sources from which we built the reference frame. Section 3 presents statistics summarising the overall properties of the reference frame in terms of the spatial distribution, accuracy, and magnitude distribution of the sources. The parallax and proper motion distributions are used as additional quality indicators and strengthen the confidence in the overall quality of the product. In Sect. 4 the optical positions in *Gaia* DR2 are compared with the VLBI frame realised in the ICRF3-prototype. A brief discussion of other quasars in the data release (Sect. 5) is followed by the conclusions in Sect. 6.

2. Construction of *Gaia*-CRF2

2.1. Principles

Starting with *Gaia* DR2, the astrometric processing of the *Gaia* data provides the parallax and the two proper motion components for most of the sources, in addition to the positions (see Lindgren 2018). As a consequence of the *Gaia* observing principle, the spin of the global reference frame must be constrained in some way in order to deliver stellar proper motions in a non-rotating frame. Less relevant for the underlying physics, but of great practical importance, is that the orientation of the resulting

Gaia frame should coincide with the current best realisation of the ICRS in the radio domain as well as possible, as implemented by the ICRF2 and soon by the ICRF3.

These two objectives were achieved in the course of the iterated astrometric solution by analysing the provisional positions and proper motions of a pre-defined set of sources, and by adjusting the source and attitude parameters accordingly by means of the so-called frame rotator (Lindgren et al. 2012). Two types of sources were used for this purpose: a few thousand sources identified as the optical counterparts of ICRF sources were used to align the positions with the radio frame, and a much larger set of probable quasars found by a cross-match with existing quasar catalogues were used, together with the ICRF sources, to ensure that the set of quasar proper motions was globally non-rotating. The resulting solution is then a physical realisation of the *Gaia* frame that is rotationally stabilised on the quasars. The detailed procedure used for *Gaia* DR2 is described in Lindgren (2018).

2.2. Selection of quasars

Although *Gaia* is meant to be autonomous in terms of the recognition of quasars from their photometric properties (colours, variability), this functionality was not yet implemented for the first few releases. Therefore the sources that are currently identified as quasars are known objects drawn from available catalogues and cross-matched to *Gaia* sources by retaining the nearest positional match. In *Gaia* DR1, quasars were flagged from a compilation made before the mission (Andrei et al. 2014), and a subset of ICRF2 was used for the alignment. The heterogeneous spatial distribution of this compilation did not greatly affect the reference frame of *Gaia* DR1 because of the special procedures that were used to link it to the HCRF (Lindgren et al. 2016; Mignard et al. 2016).

For *Gaia* DR2, which is the first release that is completely independent of the earlier HIPPARCOS and *Tycho* catalogues, it was desirable to use the most recent VLBI positions for the orientation of the reference frame, and a large, homogeneous set of quasars for the rotation. The *Gaia* data were therefore cross-matched with two different sets of known quasars:

- A prototype of the upcoming ICRF3, based on the VLBI solution of the Goddard Space Flight Center (GSFC) that comprises 4262 radio-loud quasars that are observed in the X (8.5 GHz) and S (2.3 GHz) bands. This catalogue, referred to here as the ICRF3-prototype, was kindly provided to the *Gaia* team by the IAU Working Group on ICRF3 (see Sect. 4) more than a year in advance of the scheduled release of the ICRF3. The positional accuracy is comparable to that of *Gaia*, and this set was used to align the reference frame of *Gaia* DR2 to the radio frame.
- The all-sky sample of 1.4 million active galactic nuclei (AGNs) of Secrest et al. (2015), referred to below as the AllWISE AGN catalogue (AW in labels and captions). This catalogue resulted from observations by the Wide-field Infrared Survey Explorer (WISE; Wright et al. 2010) that operates in the mid-IR at 3.4, 4.6, 12, and 22 μm wavelength. The AllWISE AGN catalogue has a relatively homogenous sky coverage, except for the Galactic plane, where the coverage is less extensive because of Galactic extinction and confusion by stars, and at the ecliptic poles, which have a higher density because of the WISE scanning law. The sources are classified as AGNs from a two-colour infrared photometric criterion, and Secrest et al. (2015) estimated that the probability of stellar contamination is $\leq 4.0 \times 10^{-5}$ per

¹ “Rotation” here refers exclusively to the kinematical rotation of the spatial axes of the barycentric celestial reference system (BCRS), as used in the *Gaia* catalogue, with respect to distant extragalactic objects (see e.g. Klioner & Soffel 1998). Similarly, “orientation” refers to the (non-) alignment of the axes at the reference epoch J2015.5.

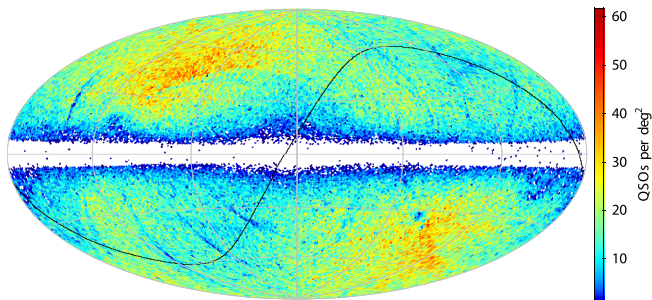


Fig. 1. Sky density per square degree for the quasars of *Gaia*-CRF2 on an equal-area Hammer–Aitoff projection in Galactic coordinates. The Galactic centre is at the origin of coordinates (centre of the map), Galactic north is up, and Galactic longitude increases from right to left. The solid black line shows the ecliptic. The higher density areas are located around the ecliptic poles.

source. About half of the AllWISE AGN sources have an optical counterpart that is detected at least once by *Gaia* in its first two years.

Cross-matching the two catalogues with a provisional *Gaia* solution and applying some filters based on the *Gaia* astrometry (see Sect. 5.1, Eq.(13) in [Lindegren 2018](#)) resulted in a list of 492 007 putative quasars, including 2844 ICRF3-prototype objects. The filters select sources with good observation records, a parallax formal uncertainty < 1 mas, a reliable level of significance in parallax and proper motion, and they avoid the Galactic plane by imposing $|\sin b| > 0.1$. These sources were used by the frame rotator, as explained above, when calculating the final solution; in *Gaia* DR2, they are identified by means of the flag `frame_rotator_object_type`. This subset of (presumed) quasars cannot, however, be regarded as a proper representation of *Gaia*-CRF2 because of the provisional nature of the solution used for the cross-matching and the relatively coarse selection criteria. Several of these sources were indeed later found to be Galactic stars.

A new selection of quasars was therefore made after *Gaia* DR2 was completed. This selection took advantage of the higher astrometric accuracy of *Gaia* DR2 and applied better selection criteria that are detailed in Sect. 5.2, Eq. (14), of [Lindegren \(2018\)](#). In particular, this updated selection takes the parallax zeropoint into account. This resulted in a set of 555 934 *Gaia* DR2 sources that are matched to the AllWISE AGN catalogue and 2820 sources that are matched to the ICRF3-prototype. The union of the two sets contains 556 869 *Gaia* DR2 sources. These sources and their positions in *Gaia* DR2 are a version of the *Gaia*-CRF that we call *Gaia*-CRF2.

The entire subsequent analysis in this paper (except in Sect. 5) is based on this sample or on subsets of it. For simplicity, we use the term quasar (QSO) for these objects, although other classifications (BL Lac object, Seyfert 1, etc.) may be more appropriate in many cases, and a very small number of them may be distant (> 1 kpc) Galactic stars.

3. Properties of *Gaia*-CRF2

This section describes the overall astrometric and photometric properties in *Gaia* DR2 of the sources of the *Gaia*-CRF2, that is, the 556 869 quasars we obtained from the match to the AllWISE AGN catalogue and the ICRF3-prototype. Their sky density is displayed in Fig. 1. The Galactic plane area is filtered out by the AllWISE AGN selection criteria, while areas around the ecliptic

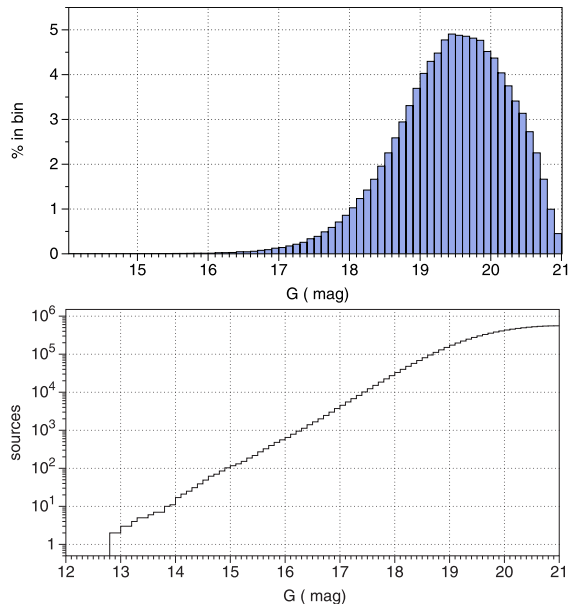


Fig. 2. G -magnitude distribution of the *Gaia*-CRF2 quasars. Percentage per bin of 0.1 mag (*top*) and cumulative distribution (*bottom*).

poles are higher than the average density, as noted above. Lower density arcs from the WISE survey are also visible, but as a rule, the whole-sky coverage outside the Galactic plane has an average density of about 14 sources per deg^2 . The few sources in the Galactic plane area come from the ICRF3-prototype.

3.1. Magnitude and colour

Figure 2 shows the magnitude distribution of the *Gaia*-CRF2 sources. In rounded numbers, there are 27 000 sources with $G < 18$ mag, 150 000 with $G < 19$ mag, and 400 000 with $G < 20$ mag. The average density of one source per square degree is reached at $G = 18.2$ mag, where it is likely that the sample is nearly complete outside the Galactic plane.

Figure 3 shows the distribution in colour index $G_{\text{BP}} - G_{\text{RP}}$ (for the definition of the blue and red passbands, BP and RP, see [Riello et al. 2018](#)). Of the sources, 2154 (0.4%) have no colour index $G_{\text{BP}} - G_{\text{RP}}$ in *Gaia* DR2. The distribution is rather narrow with a median of 0.71 mag and only 1% of the sources bluer than 0.28 mag or redder than 1.75 mag.

The magnitude is not evenly distributed on the sky, as shown in Fig. 4, with on the average fainter sources around the ecliptic poles, where the highest densities are found as well (Fig. 1). These two features result from a combination of the deeper survey of AllWISE in these areas and the more frequent *Gaia* observations from the scanning law.

3.2. Astrometric quality

In this section we describe the astrometric quality of the *Gaia*-CRF2 quasars based on the formal positional uncertainties and on the distribution of observed parallaxes and proper motions, which are not expected to be measurable by *Gaia* at the level of individual sources. We defer a direct comparison of the *Gaia* positions with VLBI astrometry to Sect. 4.

3.2.1. Formal uncertainty in position

As a single number characterising the positional uncertainty of a source, $\sigma_{\text{pos,max}}$, we take the semi-major axis of the dispersion

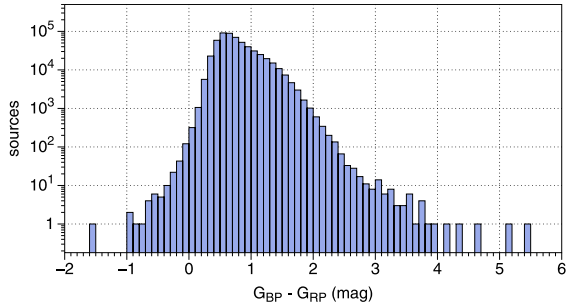


Fig. 3. Colour distribution of the *Gaia*-CRF2 quasars (log scale for the number of sources per bin of 0.1 mag).

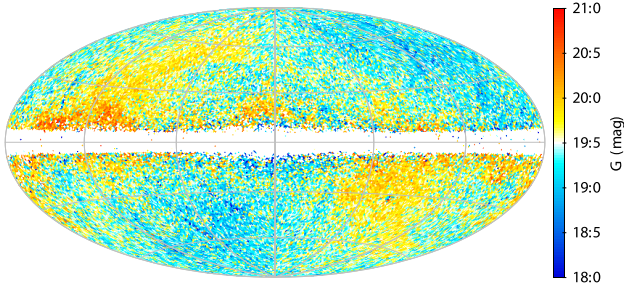


Fig. 4. Sky distribution of the *Gaia*-CRF2 source magnitudes. This map shows the median values of the G magnitude in cells of approximately 0.84 deg^2 using an equal-area Hammer–Aitoff projection in Galactic coordinates. The Galactic centre is at the origin of coordinates (centre of the map), Galactic north is up, and Galactic longitude increases from right to left.

ellipse, computed from a combination $\sigma_{\alpha^*} = \sigma_{\alpha} \cos \delta$, σ_{δ} , and the correlation coefficient $\varrho_{\alpha,\delta}$:

$$\sigma_{\text{pos,max}}^2 = \frac{1}{2} \left(\sigma_{\alpha^*}^2 + \sigma_{\delta}^2 + \sqrt{(\sigma_{\alpha^*}^2 - \sigma_{\delta}^2)^2 + (2\sigma_{\alpha^*}\sigma_{\delta}\varrho_{\alpha,\delta})^2} \right). \quad (1)$$

Because this is also the highest eigenvalue of the 2×2 covariance matrix, it is invariant to a change of coordinates. These are formal uncertainties (see Sect. 3.2.2 for a discussion of how real they are) for the reference epoch J2015.5 of *Gaia* DR2. The results for the whole sample of *Gaia*-CRF2 quasars and the subset with $G < 19$ are shown in Fig. 5. The median accuracy is 0.40 mas for the full set and 0.20 mas for the brighter subset. Additional statistics are given in Table 1.

The main factors governing the positional accuracy are the magnitude (Fig. 6) and location on the sky (Fig. 7). The larger-than-average uncertainty along the ecliptic in Fig. 7 is conspicuous; this is a signature of the *Gaia* scanning law. This feature will also be present in future releases of *Gaia* astrometry and will remain an important characteristic of the *Gaia*-CRF.

3.2.2. Parallaxes and proper motions

Parallaxes and proper motions are nominally zero for the quasars that were selected for the reference frame (we neglect here the expected global pattern from the Galactic acceleration, which is expected to have an amplitude of $4.5 \mu\text{as yr}^{-1}$, see Sect. 3.3). Their statistics are useful as complementary indicators of the global quality of the frame and support the accuracy claim. Here we consider the global statistics before investigating possible systematics in Sect. 3.3. Figure 8 shows the distribution of the parallaxes for different magnitude-limited subsets. As explained

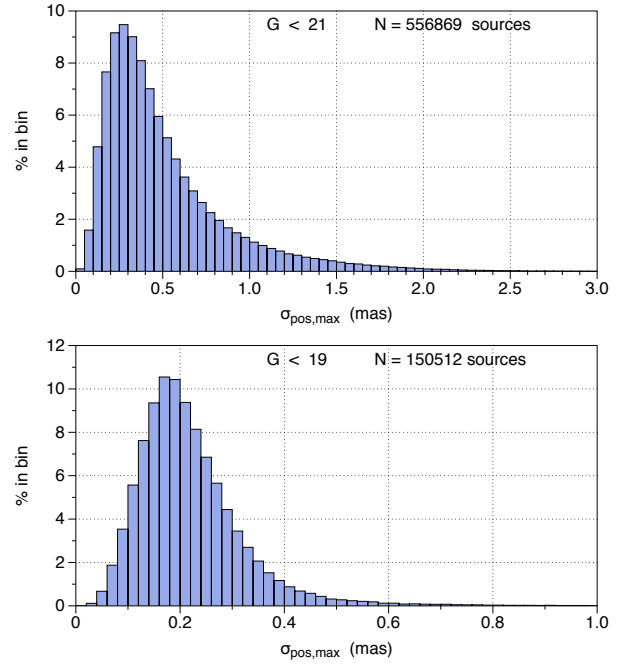


Fig. 5. Distribution of positional uncertainties $\sigma_{\text{pos,max}}$ for the *Gaia*-CRF2 quasars: all sources (top) and $G < 19$ mag (bottom).

Table 1. Positional uncertainty $\sigma_{\text{pos,max}}$ of the *Gaia*-CRF2 quasars.

G selection [mag]	N	1st quartile [mas]	median [mas]	3rd quartile [mas]
< 18.0	27 275	0.09	0.12	0.15
$[18.0-19.0]$	123 237	0.17	0.22	0.28
$[19.0-19.5]$	125 029	0.27	0.33	0.41
$[19.5-20.0]$	132 418	0.38	0.47	0.59
≥ 20.0	148 910	0.61	0.81	1.12
all	556 869	0.26	0.40	0.64

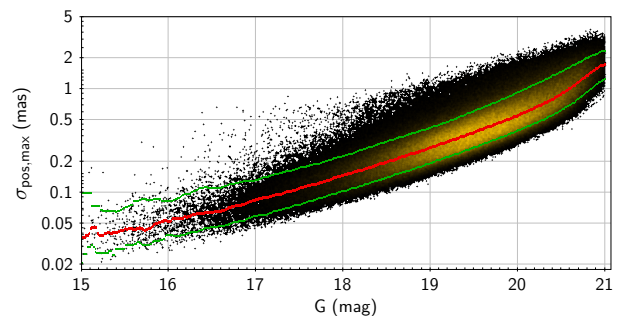


Fig. 6. Positional uncertainties $\sigma_{\text{pos,max}}$ for the sources in the *Gaia*-CRF2 as function of the G magnitude. The red solid line is the running median, and the two green lines are the first and ninth decile.

in Lindegren (2018), the *Gaia* DR2 parallaxes have a global zero-point error of -0.029 mas, which is not corrected for in the data available in the *Gaia* archive. This feature is well visible for the quasar sample and is a real instrumental effect that is not yet eliminated by the calibration models. Fortunately, the offset is similar for the different subsets. The shape of the distributions (best visible in the full set) is clearly non-Gaussian because of the

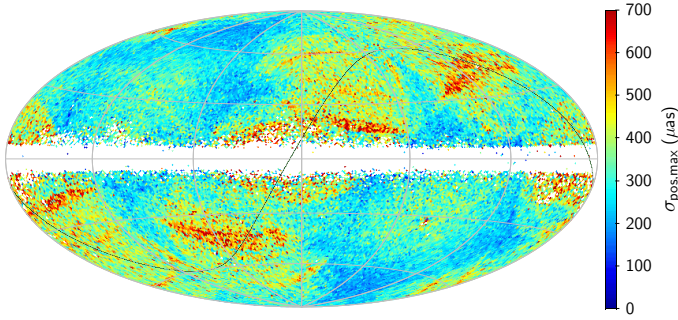


Fig. 7. Spatial distribution of the formal position uncertainty in Eq. (1) for the 407 959 sources of the *Gaia*-CRF2 with $G < 20$. The map shows the median value in each cell of approximately 0.84 deg^2 , using a Hammer–Aitoff projection in Galactic coordinates with zero longitude at the centre and increasing longitude from right to left. The solid black line shows the ecliptic.

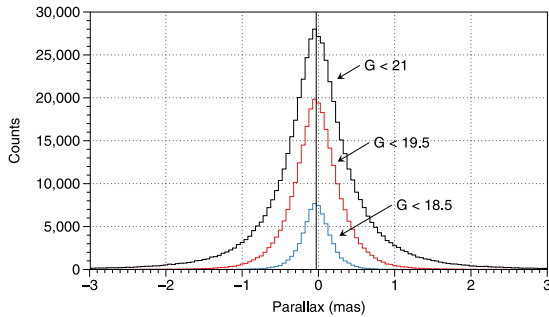


Fig. 8. Distribution of parallaxes in the *Gaia* archive for the *Gaia*-CRF2 quasars, subdivided by the maximum magnitude. The line at $\varpi = -0.029 \text{ mas}$ shows the global zeropoint offset.

mixture of normal distributions with a large spread in standard deviation, which is primarily linked to the source magnitude. The typical half-widths of the distributions (0.4, 0.3, and 0.2 mas) are of a similar size as the median positional uncertainties in Table 1.

The distribution of the normalised debiased parallaxes, computed as $(\varpi + 0.029 \text{ mas})/\sigma_{\varpi}$, should follow a standard normal distribution (zero mean and unit variance) if the errors are Gaussian and the formal uncertainties σ_{ϖ} are correctly estimated. The actual distribution for the full set of 556 869 quasars is plotted in Fig. 9. The red continuous curve is a normal distribution with zero mean and standard deviation 1.08; that this very closely follows the distribution up to normalised values of ± 3.5 is an amazing feature for real data. The magnitude effect is then fully absorbed by the normalisation, indicating that the *Gaia* accuracy in this brightness range is dominated by the photon noise. The factor 1.08 means that the formal uncertainties of the parallaxes are too small by 8%.

Similarly, the distributions in Fig. 10 for the normalised components of proper motions are very close to a normal distribution, with zero mean and standard deviations of 1.09 (μ_{α^*}) and 1.11 (μ_{δ}). The extended distributions in log scale are very similar to the parallax and are not plotted.

3.3. Systematic effects

3.3.1. Spatial distributions

In an ideal world, the errors in position, parallax, and proper motion should be purely random and not display any systematic

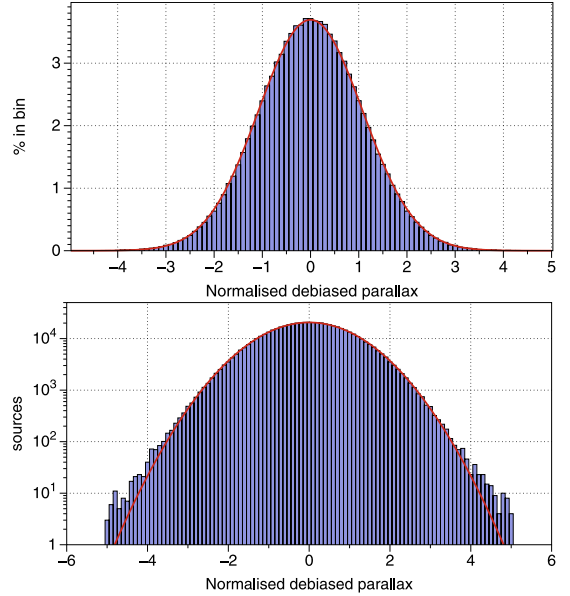


Fig. 9. Distribution of the normalised debiased parallaxes, $(\varpi + 0.029 \text{ mas})/\sigma_{\varpi}$, for the *Gaia*-CRF2 quasars in linear scale (*top*) and logarithmic (*bottom*). The red curve is a normal distribution with zero mean and standard deviation 1.08.

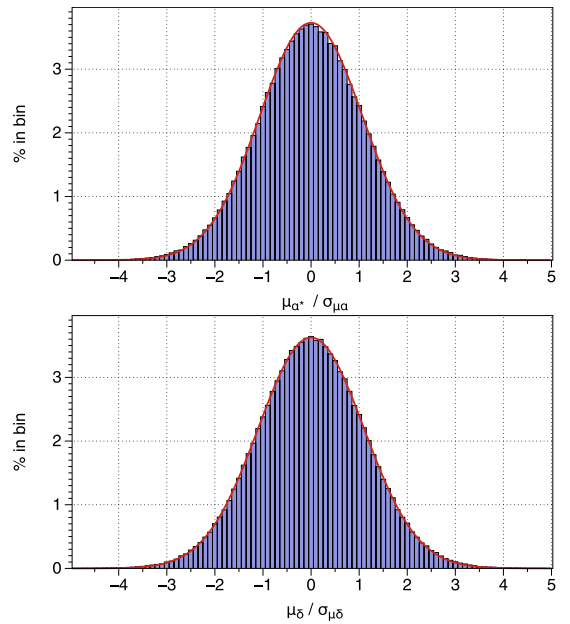


Fig. 10. Distributions of the normalised components of proper motions of the QSOs found with *Gaia* data, with μ_{α^*} (*top*) and μ_{δ} (*bottom*). A normal distribution with zero mean and standard deviation of 1.09 for μ_{α^*} (1.11 for μ_{δ}) is drawn in red.

patterns as function of position on the celestial sphere. While the non-uniform sampling of the sky produced by the *Gaia* scanning law is reflected in the formal uncertainties of the quasar astrometry, as shown in Fig. 7 for the positions, this does not imply that the errors (i.e. the deviations from the true values) show patterns of a similar nature. In the absence of a reliable external reference for the positions (except for the VLBI subset), the possibility of investigating the true errors in position is limited. However, the positions are derived from the same set of observations as the other astrometric parameters, using the same solution. Since the errors in parallax σ and proper motion are found

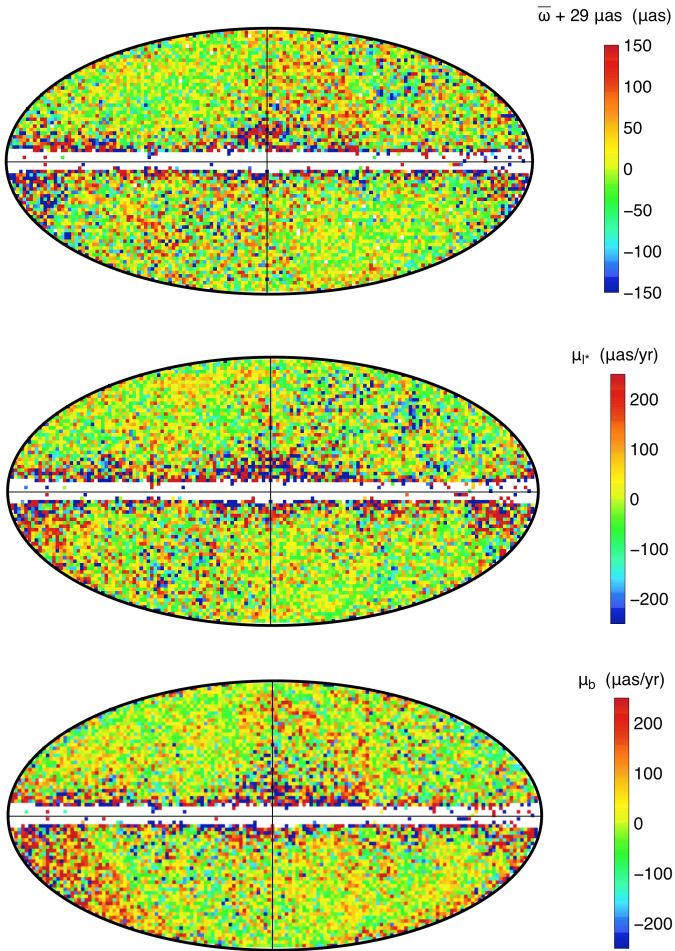


Fig. 11. Spatial distributions (in Galactic coordinates) of the parallaxes and proper motions of the *Gaia*-CRF2 quasars. From *top to bottom*: parallax ($\bar{\omega}$), proper motion in Galactic longitude (μ_{l^*}), and proper motion in Galactic latitude (μ_b). Median values are computed in cells of approximately 5 deg^2 . Maps are in the Hammer–Aitoff projection with Galactic longitude zero at the centre and increasing from right to left.

to be in good agreement with the formal uncertainties calculated from the solution, we expect this to be the case for the positional errors as well.

Figure 11 shows maps in Galactic coordinates of the median parallax and proper motion components of the *Gaia*-CRF2 sources, calculated over cells of 4.669 deg^2 . For cells of this size, the median number of sources per cell is 70, with the exception of low Galactic latitude, where the density is lower (see Fig. 1), resulting in a larger scatter of the median from cell to cell than in other parts of the map. In Fig. 11 this is visible as an increased number of cells with red and blue colours, instead of green and yellow, in the less populated areas.

The median parallaxes shown in the top panel of Fig. 11 were corrected for the global zeropoint of $-29 \mu\text{as}$ (Sect. 3.2.2). In all three maps, various large-scale patterns are seen for Galactic latitudes $|b| \gtrsim 10\text{--}15 \text{ deg}$, while at small angles (cell size), only a mixture of positive or negative offsets is visible that results from normal statistical scatter. The visual interpretation is complicated by large-scale patterns in the amplitude of the statistical scatter, in particular the smaller scatter in the second and fourth quadrants, that is, around the ecliptic poles. This is the result of a combination of the sky distribution of the sources

(Fig. 1), their magnitudes (Fig. 4), and the *Gaia* scanning law (Fig. 7), which all exhibit similar patterns. Quantifying the large-scale systematics therefore requires a more detailed numerical analysis.

3.3.2. Spectral analysis

The vector field of the proper motions of the *Gaia*-CRF2 quasars was analysed using expansions on a set of vector spherical harmonics (VSH), as explained in Mignard & Klioner (2012) or Vityazev & Tsvetkov (2014).

In this approach the components of proper motion are projected onto a set of orthogonal functions up to a certain degree l_{max} . The terms of lower degrees provide global signatures such as the rotation and other important physical effects (secular acceleration, gravitational wave signatures), while harmonics of higher degree hold information on local distortions at different scales. Given the patterns seen in Fig. 11, we expect to see a slow decrease in the power of harmonics with $l > 1$. The harmonics of degree $l = 1$ play a special role, since any global rotation of the system of proper motions will be observed in the form of a rotation vector directly extracted from the three components with $(l, m) = (1, 0)$, $(1, -1)$, and $(1, +1)$, where m is the order of the harmonic ($|m| \leq l$).

Mignard & Klioner (2012) derived a second global term from $l = 1$ that they called glide. This physically corresponds to a dipolar displacement originating at one point on a sphere and ending at the diametrically opposite point. For the quasar proper motions, this vector field is precisely the expected signature of the galactocentric acceleration (Fanselow 1983; Bastian 1995; Sovers et al. 1998; Kovalevsky 2003; Titov & Lambert 2013). As summarised in Table 2, several VSH fits were made using different selections of quasars or other configuration parameters. Fit 1 uses all the quasars and fits only the rotation, without glide or harmonics with $l > 1$. This is very close to the conditions used to achieve the non-rotating frame in the astrometric solution for *Gaia* DR2. It is therefore not surprising that the rotation we find is much smaller than in the other experiments. The remaining rotation can be explained by differences in the set of sources used, treatment of outliers, and so on. This also illustrates the difficulty of producing a non-rotating frame that is non-rotating for every reasonable subset that a user may wish to select: This is not possible, at least at the level of formal uncertainties. Experiment 2 fits both the rotation and glide to all the data. The very small change in rotation compared with fit 1 shows the stability of the rotation resulting from the regular spatial distribution of the sources and the consequent near-orthogonality of the rotation and glide on this set. Fit 3 includes all harmonics of degree $l \leq 5$, that is, 70 fitted parameters. Again the results do not change very much because of the good spatial distribution. The next five fits show the influence of the selection in magnitude and modulus of proper motion, and of not weighting the data by the inverse formal variance. In the next two fits (9a and 9b), the data are divided into two independent subsets, illustrating the statistical uncertainties. Most of these fits use fewer sources with a less regular distribution on the sky.

The last fit, fit 10, uses only the faint sources and has a similar glide but a very different rotation (x and y components, primarily), although it comprises the majority (73%) of the *Gaia*-CRF2 sources. This agrees with Figs. 3 and 4 in Lindegren (2018), which show a slight dependency on colour and magnitude of the *Gaia* spin relative to quasars. Again, this illustrates the

Table 2. Large-scale structure of the proper motion field of the *Gaia*-CRF2 quasars analysed using vector spherical harmonics.

Fit	Source selection	W	l_{\max}	N	Rotation [$\mu\text{as yr}^{-1}$]			Glide [$\mu\text{as yr}^{-1}$]		
					x	y	z	x	y	z
1	all	y	1	556869	-3.1 ± 0.8	-1.9 ± 0.7	-1.0 ± 0.9	–	–	–
2	all	y	1	556869	-3.6 ± 0.8	-2.2 ± 0.7	-0.9 ± 0.9	-7.0 ± 0.8	4.7 ± 0.7	12.1 ± 0.7
3	all	y	5	556869	-5.5 ± 1.1	-7.4 ± 0.9	5.6 ± 1.2	-9.2 ± 1.2	4.7 ± 1.0	11.6 ± 1.0
4	$\mu < 2 \text{ mas yr}^{-1}, G < 18$	y	5	27189	-13.8 ± 2.0	-13.2 ± 1.7	4.0 ± 2.2	-7.9 ± 2.2	4.7 ± 1.8	10.3 ± 1.7
5	$\mu < 2 \text{ mas yr}^{-1}, G < 18$	n	5	27189	-8.9 ± 2.9	-12.1 ± 2.4	2.8 ± 2.5	-10.4 ± 2.9	5.7 ± 2.4	16.6 ± 2.5
6	$\mu < 2 \text{ mas yr}^{-1}, G < 19$	y	5	149146	-11.2 ± 1.3	-12.0 ± 1.1	4.4 ± 1.4	-9.8 ± 1.5	4.6 ± 1.2	10.4 ± 1.1
7	$\mu < 3 \text{ mas yr}^{-1}, G < 20$	y	5	400472	-5.9 ± 1.1	-8.6 ± 0.9	5.1 ± 1.2	-9.0 ± 1.2	4.1 ± 1.0	11.9 ± 1.0
8	$\mu < 3 \text{ mas yr}^{-1}$	y	5	513270	-5.7 ± 1.1	-7.9 ± 0.9	5.2 ± 1.2	-8.8 ± 1.2	4.1 ± 1.0	11.6 ± 0.9
9a	$[10^5 \alpha] \bmod 2 = 0$	y	5	278170	-5.8 ± 1.6	-8.9 ± 1.3	6.4 ± 1.7	-8.5 ± 1.7	3.0 ± 1.4	12.5 ± 1.4
9b	$[10^5 \alpha] \bmod 2 = 1$	y	5	278699	-5.1 ± 1.6	-5.8 ± 1.3	4.8 ± 1.7	-9.8 ± 1.7	6.6 ± 1.4	10.7 ± 1.4
10	$G > 19$	y	5	406356	9.8 ± 2.1	6.2 ± 1.8	7.0 ± 2.4	-8.3 ± 2.3	3.3 ± 1.9	15.6 ± 1.9

Notes. $\mu = (\mu_{\alpha^*}^2 + \mu_{\delta}^2)^{1/2}$ is the modulus of the proper motion. N is the number of sources used in the solution. $W = \text{“y”}$ or “n” for weighted or unweighted solution. The weighted solutions use a block-diagonal weight matrix obtained from the 2×2 covariance matrix of each source. l_{\max} is the highest degree of the fitted VSH from which rotation and glide are extracted for $l = 1$. The columns headed x, y, z give the components of the rotation and glide along the principal axes of the ICRS. In rows 9a and 9b, two independent halves of the sample are selected according to whether $[10^5 \alpha]$ is even (9a) or odd (9b), with α in degrees.

sensitivity of the determination of the residual spin to the source selection, and at this stage, we cannot offer a better explanation than that a single solid rotation is too simple a model to fit the entire range of magnitudes. No attempt was made to introduce a magnitude equation in the fits.

The formal uncertainty of all the fits using at least a few hundred thousand quasars is of the order of $1 \mu\text{as yr}^{-1}$. It is tempting to conclude from this that the residual rotation of the frame with respect to the distant universe is of a similar magnitude. However, the scatter from one fit to the next is considerably larger, with some values exceeding $10 \mu\text{as yr}^{-1}$. Clearly, an overall solid rotation does not easily fit all the *Gaia* data, but gives results that vary with source selection well above the statistical noise. However, the degree of consistency between the various selections allows us to state that the residual rotation rate of the *Gaia*-CRF is probably not much higher than $\pm 10 \mu\text{as yr}^{-1}$ in each axis for any subset of sources.

The typical glide vector is about $(-8, +5, +12) \pm 1 \mu\text{as yr}^{-1}$ for the components in the ICRS. The expected signature for the galactocentric acceleration is a vector directed towards the Galactic centre with a magnitude of $\approx 4.50 \mu\text{as yr}^{-1}$, or $(-0.25, -3.93, -2.18) \mu\text{as yr}^{-1}$ in the ICRS components. Clearly, the large-scale systematic effects in the *Gaia* proper motions, being of the order of $10 \mu\text{as yr}^{-1}$ at this stage of the data analysis, prevent a fruitful analysis of the quasar proper motion field in terms of the Galactic acceleration. For this purpose, an order-of-magnitude improvement is needed in the level of systematic errors, which may be achieved in future releases of *Gaia* data based on better instrument calibrations and a longer observation time-span. A similar improvement is needed to achieve the expected estimate of the energy flux of the primordial gravitational waves (Gwinn et al. 1997; Mignard & Klioner 2012; Klioner 2018).

The overall stability of the fits in Table 2 is partly due to the fairly uniform distribution of the *Gaia*-CRF2 sources over the celestial sphere, and it does not preclude the existence of significant large-scale distortions of the system of proper

motions. Such systematics may be quantified by means of the fitted VSH, however, and a convenient synthetic indicator of how much signal is found at different angular scales is given by the total power in each degree l of the VSH expansion. This power \mathcal{P}_l is invariant under orthogonal transformation (change of coordinate system) and therefore describes a more intrinsic, geometric feature than the individual components of the VSH expansion. The degree l corresponds to an angular scale of $\sim 180^\circ/l$.

In Fig. 12 (top panel) we plot $(\mathcal{P}_l/4\pi)^{1/2}$ in $\mu\text{as yr}^{-1}$, representing the RMS value of the vector field for the corresponding degree l . The lower panel in Fig. 12 shows the significance level of the power given as the equivalent standard normal variate derived from the asymptotic χ^2 distribution; see Mignard & Klioner (2012) for details. The points labelled S and T correspond to the spheroidal and toroidal harmonics, with T & S for their quadratic combination. To illustrate the interpretation of the diagrams, for T_1 the RMS value is $(\mathcal{P}_1/4\pi)^{1/2} \approx 10 \mu\text{as yr}^{-1}$, which should be similar to the magnitude of the rotation vector for fit 3 in Table 2. The significance of this value is $Z_{\chi^2} \approx 7$, corresponding to 7σ of a normal distribution, or a probability below 10^{-11} .

For the low degrees plotted in Fig. 12, the power generally decreases with increasing l (smaller angular scales). This indicates that the systematics are generally dominated by the large angular scales. The total RMS for $l \leq 10$ (angular scales $\gtrsim 18 \text{ deg}$) is $42 \mu\text{as yr}^{-1}$.

Lindgren (2018) analysed the large-scale systematics of the *Gaia* DR2 proper motions of exactly the same quasar sample, using a very different spatial correlation technique. A characteristic angular scale of 20 deg was found, with an RMS amplitude of $28 \mu\text{as yr}^{-1}$ per component of proper motion (their Eq. (18)). Since this corresponds to $40 \mu\text{as yr}^{-1}$ for the total proper motion, their result is in good agreement with ours. They also found higher-amplitude oscillations with a spatial period of $\approx 1 \text{ deg}$, which in the present context of *Gaia*-CRF2 are almost indistinguishable from random noise, however.

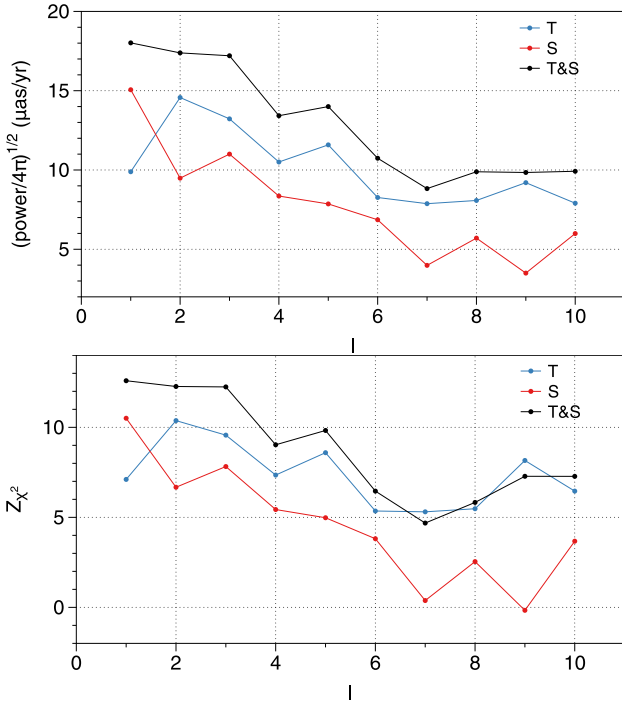


Fig. 12. Distributions of the RMS values (*top*) and their statistical levels of significance (*bottom*) in the VSH decomposition of the proper motion vector field of the *Gaia*-CRF2 from $l = 1$ to 10. *S* and *T* refer to the spheroidal and toroidal harmonics, and *T&S* signifies their quadratic combination.

4. ICRF3-prototype subset of *Gaia*-CRF2

This section describes the subset of 2820 *Gaia*-CRF2 quasars matched to the ICRF3-prototype (Sect. 2.2), that is, the optical counterparts of compact radio sources with accurate VLBI positions. A comparison between the optical and VLBI positions is in fact a two-way exercise, as useful for understanding the radio frame as it is to *Gaia*, since neither of the two datasets is significantly better than the other. A similar investigation of the reference frame for *Gaia* DR1 (Mignard et al. 2016) showed the limitations of ICRF2, the currently available realisation of the ICRS, for such a comparison. A subset ICRF2 sources also had a less extensive VLBI observation record, the accuracy was lower for the best sources, and it would have been only marginally useful for a comparison to the *Gaia* DR2.

In discussions with the IAU working group in charge of preparing the upcoming ICRF3, which is scheduled for mid-2018, it was agreed that the working group would provide a prototype version of ICRF3 in the form of their best current solution to the *Gaia* team. This ICRF3-prototype was officially delivered in July 2017 and is particularly relevant in the current context for two reasons.

- With the assumption that there is no globally systematic difference between the radio and optical positions, the common sources allowed the axes of the two reference frames to be aligned, as explained in Lindegren (2018). The existence of radio–optical offsets with random orientation for each source is not a great problem for this purpose as it only adds white noise to the position differences. If large enough, it will be detected in the normalised position differences (Sect. 4.3).
- The VLBI sources included in this prototype, together with the associated sets worked out in the X/Ka and *K* band

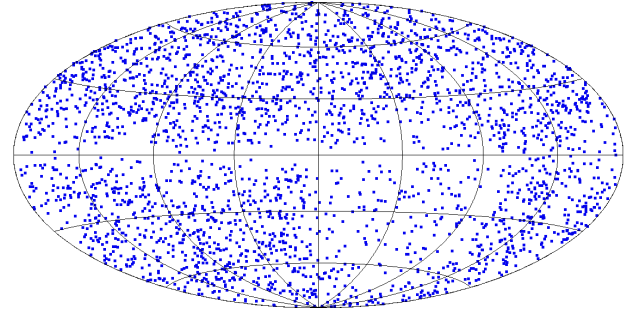


Fig. 13. Sky distribution of the 2820 *Gaia* sources identified as most probable optical counterparts of quasars in the ICRF3-prototype. Hammer–Aitoff projection in Galactic coordinates with origin at the centre of the map and longitude increasing from right to left.

(not yet released), are the most accurate global astrometric solutions available today that are fully independent of *Gaia*. The quoted uncertainties are very similar to what is formally achieved in *Gaia* DR2, and the best-observed VLBI sources have positions that are nominally better than those from *Gaia*. This is therefore the only dataset from which the true errors and possible systematics in the positions of either dataset can be assessed and individual cases of truly discrepant positions between the radio and optical domains can be identified. The VLBI positions are less homogeneous in accuracy than the corresponding *Gaia* data, but the ≈ 1650 ICRF3-prototype sources with a (formal) position uncertainty < 0.2 mas match the *Gaia* positions of the brighter ($G < 18$ mag) sources well in quality.

4.1. Properties of the *Gaia* sources in the ICRF3-prototype

Figure 13 shows the spatial distribution of the 2820 optical counterparts of ICRF3-prototype sources on the sky. The plot is in Galactic coordinates to facilitate comparison with Fig. 1, showing the full *Gaia*-CRF2 sample. The area in the lower right quadrant with low density corresponds to the region of the sky at $\delta < -40$ deg with less VLBI coverage. Otherwise the distribution is relatively uniform, but with a slight depletion along the Galactic plane, as expected for an instrument operating at optical wavelengths.

The magnitude distribution of the ICRF3-prototype sources is shown in Fig. 14. The median is 18.8 mag, compared with 19.5 mag for the full *Gaia*-CRF2 sample shown in Fig. 2. The colour distribution (not shown) is similar to that of the full sample, shown in Fig. 2, only slightly redder: the median $G_{BP} - G_{RP}$ is ≈ 0.8 mag for the ICRF3-prototype subset, compared with 0.7 mag for the full sample.

In terms of astrometric quality, the *Gaia* DR2 sources in the ICRF3-prototype subset do not differ significantly from other quasars in *Gaia*-CRF2 at the same magnitude. Figure 15 displays the formal uncertainty in position, computed with Eq. (1), as function of the *G* magnitude. Both the median relation and the scatter about the median are virtually the same as for the general population of quasars in *Gaia*-CRF2 shown in Fig. 6. For $G \gtrsim 16.2$, only few points in Fig. 15 lie clearly above the main relation. This may be linked to the change in the onboard CCD observation window allocation that occurs at $G \approx 16$ (Gaia Collaboration et al. 2016). Four hundred and nine sources are brighter than $G = 17.4$, where the median position uncertainty as shown on Fig. 15 reaches 100 μ as.

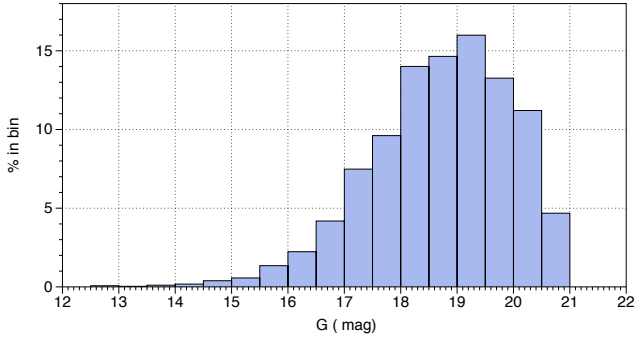


Fig. 14. Magnitude distribution of the 2820 *Gaia* sources identified as likely optical counterparts of quasars in the ICRF3-prototype.

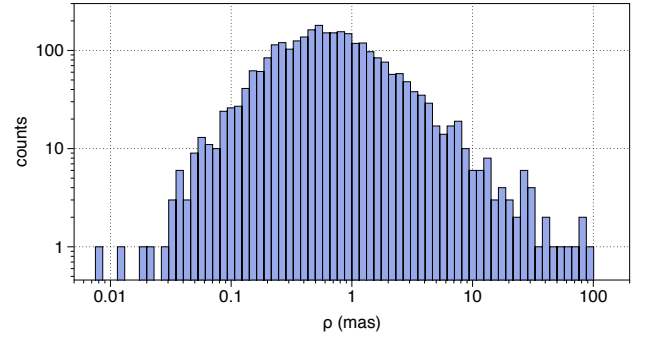


Fig. 16. Distribution of the angular separation ρ between *Gaia* DR2 and the ICRF3-prototype for the 2820 sources in common. Log-log scale plot with bins in ρ having a fractional width of $2^{1/5}$.

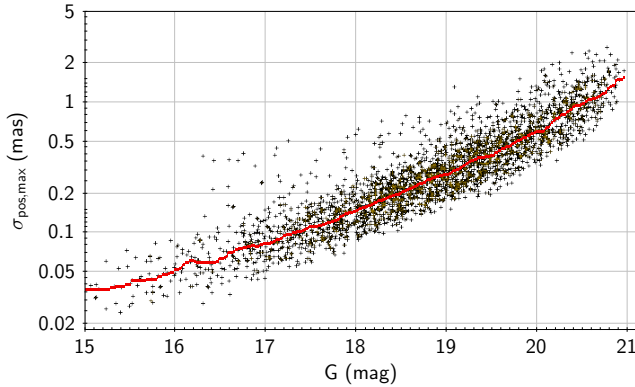


Fig. 15. Formal position uncertainty as a function of magnitude for the 2820 *Gaia* sources identified as optical counterparts of quasars in the ICRF3-prototype. The solid red line is a running median through the data points.

4.2. Angular separations

We now compare the positions in *Gaia* DR2 and ICRF3-prototype directly for the 2820 quasars in common. Figure 16 gives in log-scale the distribution of the angular distances computed as

$$\rho = (\Delta\alpha_*^2 + \Delta\delta^2)^{1/2}, \quad (2)$$

where $\Delta\alpha_* = (\alpha_{Gaia} - \alpha_{VLBI}) \cos \delta$. While for most of the sources, ρ is lower than 1 mas and very often much below this level, the number of discrepant sources is significant, and a few even have a position difference higher than 10 mas that would require individual examination.

To illustrate the dependence on the solution accuracies, Fig. 17 shows scatter plots of ρ versus the formal uncertainty in the ICRF3-prototype (top) and *Gaia*-CRF2 (bottom). Several of the most extreme distances in the top diagram are for sources with a large uncertainty in the ICRF3-prototype. However, some sources with nominally good solutions in both datasets exhibit large positional differences. These deserve more attention as the differences could represent real offsets between the centres of emission at optical and radio wavelengths. This is not further investigated in this paper, which is devoted to present the main properties of the *Gaia*-CRF2. Other explanations for the large differences can be put forward, such as a mismatch on the *Gaia* side when the optical counterpart is too faint and a distant star happens to be matched instead (unlikely at <10 mas distance); an extended galaxy around the quasar that is misinterpreted by

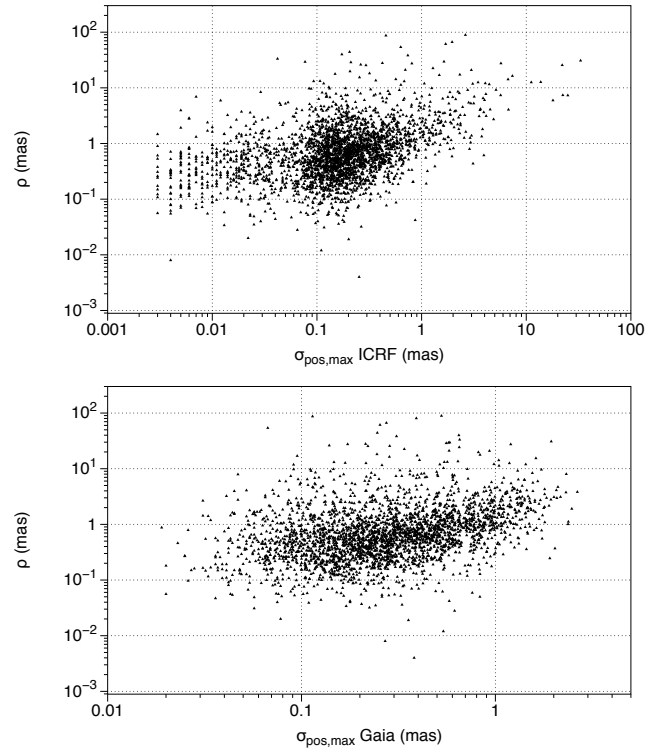


Fig. 17. Angular position differences ρ between *Gaia* DR2 and the ICRF3-prototype as function of the formal uncertainties $\sigma_{\text{pos,max}}$ of the ICRF3-prototype (top) and *Gaia* DR2 (bottom).

the *Gaia* detector (should in general produce a poor solution); double or lensed quasars; or simply statistical outliers from the possibly extended tails of random errors. Although the ICRF3-prototype data in Fig. 17 cover a wider range in $\sigma_{\text{pos,max}}$ than the *Gaia* data, the cores of both distributions extend from ≈ 0.1 to 0.5 mas.

4.3. Normalised separations

The angular separations ρ become statistically more meaningful when scaled with the combined standard uncertainties. In the case of correlated variables, Mignard et al. (2016) have shown how to compute a dimensionless statistic X , called the normalised separation (their Eq. (4)). If the positional errors in both catalogues are Gaussian with the given covariances, then X is expected to follow a standard Rayleigh distribution, and values

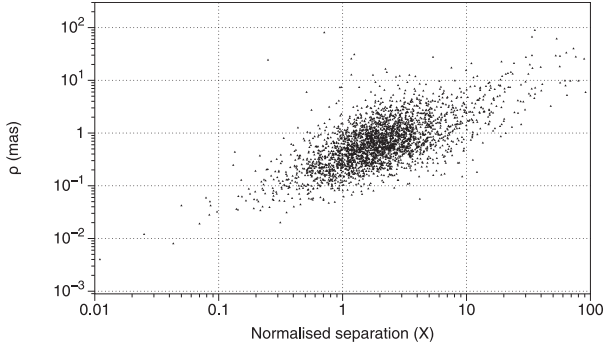


Fig. 18. Angular separation (ρ) vs. normalised separation (X) for the ICRF3-prototype subset of *Gaia*-CRF2.

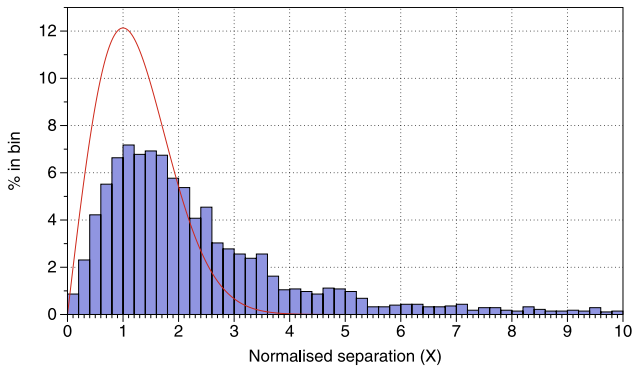


Fig. 19. Distribution of the normalised separations X between *Gaia* DR2 and the ICRF3-prototype. 148 sources have a normalised separation >10 . The red curves show the (standard) Rayleigh distributions for unit standard deviation.

of $X > 3$ should be rare (probability ≈ 0.01). We caution that the normalisation used in this section depends on the reliability of the reported position uncertainties from the *Gaia* DR2 on one hand and from the ICRF3-prototype on the other hand. The latter are still provisional and purely formal, without noise floor and other overall adjustment, which will be introduced in the final release of the ICRF3.

Figure 18 is a scatter plot of ρ versus X , showing a fairly large subset of sources with $X > 3$ and even much larger. The most anomalous cases are found in the upper right part of the diagram. Some of the sources with the largest ρ are located in the upper centre of the diagram, with unremarkable X , meaning that their large angular separations are not significant in view of the formal uncertainties.

The diagram in Fig. 19 shows the frequency distribution of the normalised separations X with the standard Rayleigh probability density function superimposed as a solid red line. The frequency diagram includes all the sources, although 148 sources have a normalised separation >10 and would be outside the frame. The distribution cannot be represented by a standard Rayleigh distribution, even though its mode is not very far from one, but the spread at large normalised separations is much too large. The departure from a pure Rayleigh distribution between VLBI positions and the *Gaia* DR1 has previously been noted in Petrov & Kovalev 2017 in a comparison using more than 6000 sources with VLBI positions. However, at this stage with the ICRF3-prototype, we cannot draw definite conclusions, and this issue will have to be reconsidered with the official release of ICRF3.

4.4. Large-scale systematics

In this section we analyse the positional difference between *Gaia* DR2 and the ICRF3-prototype in terms of large-scale spatial patterns. As in Sect. 3.3, the vector field of position differences is decomposed using VSH, where in particular the coefficients for degree $l = 1$ give the orientation difference of the two frames and a glide in position. Several fits were made to assess the stability of the orientation rotation against various selections of sources. Nominally, *Gaia* DR2 has been aligned to the ICRF3-prototype and no significant orientation difference should remain. However, stating that the two frames have been aligned is not the complete story, since the final alignment depends on many details of the fit: weighting scheme, outlier filtering, magnitude selection, and the model used for the fit. Furthermore, as explained in Sect. 2.2, the alignment was made using a slightly different set of ICRF3-prototype sources than currently considered. As a consequence of these differences, we often find statistically significant non-zero orientation errors in our fits. The amplitude of these errors provides the best answer to the question of how precisely the two frames share the same axes.

The results of the various fits are summarised in Table 3. The first fit is similar to the alignment procedure in the astrometric solution for *Gaia* DR2 in that only the three orientation parameters (otherwise denoted ϵ_x , ϵ_y , ϵ_z) are fitted without a glide component. Of all the fits in the table, this has the overall smallest, statistically most insignificant orientation parameters. It gives a formal uncertainty in the alignment of about $30 \mu\text{as}$ per axis. Fit 2, using the same data set, but fitting the glide as well, reveals a different picture. The orientation parameters remain negligible, but not as close to zero as in fit 1, and the glide components have a significant amplitude. The uncertainty is unchanged at about $30 \mu\text{as}$. This is a good illustration of the ambiguity in the alignment when the procedure is not fully implemented.

In fits 3 to 5, only the orientation parameters are estimated, but with different filtering of the data, with or without statistical weighting of the differences. We showed in Sect. 4.1 that a subset of sources has good astrometric quality in both catalogues, but the position differences are not compatible with the formal uncertainties. Removing these sources from the fit greatly improves the formal precision of the fit, while the orientation parameters are changed by a few tens of μas , which is still only marginally significant. More significant changes result from including the glide and higher degrees of VSH (fits 6 to 8), or restricting the sample to the brighter subset (fits 9 and 10) with or without weighting. In these fits particularly the orientation error in x and the glide in y become significant. Finally, cases 11a and b are run on two independent halves of the data to ascertain the sensitivity of the solution to the selection.

Based on these experiments, we state that the axes of the *Gaia*-CRF2 and the ICRF3-prototype are aligned with an uncertainty of 20 to $30 \mu\text{as}$, but no precise value can be provided without agreeing on the detailed model and numerical procedures for determining the orientation errors.

5. Other quasars in *Gaia* DR2

The cross-match of *Gaia* DR2 with the AllWISE AGN catalogue provided a very clean and homogeneous sample of quasars that is suitable for the definition of the *Gaia*-CRF2 and systematic investigation of its properties. However, other catalogues exist that will enlarge the sample of known or probable quasars in *Gaia* DR2 for other purposes. The Million

Table 3. Global differences between the *Gaia*-CRF2 positions of ICRF sources and their positions in the ICRF3-prototype, expressed by the orientation and glide parameters.

Fit	Source selection	W	l_{\max}	N	Orientation (μas)			Glide (μas)		
					x	y	z	x	y	z
1	all	y	1	2820	-9 ± 29	4 ± 27	3 ± 28	–	–	–
2	all	y	1	2820	-28 ± 31	-8 ± 29	10 ± 28	47 ± 29	-69 ± 28	-72 ± 29
3	$\rho < 10$ mas	y	1	2773	-17 ± 16	22 ± 15	-23 ± 16	–	–	–
4	$\rho < 2$ mas	y	1	2423	-35 ± 9	21 ± 8	-24 ± 9	–	–	–
5	$\rho < 2$ mas	n	1	2423	-13 ± 14	5 ± 14	-5 ± 13	–	–	–
6	$\rho < 2$ mas	y	5	2423	-47 ± 12	30 ± 10	0 ± 11	2 ± 12	-40 ± 10	-25 ± 11
7	$\rho < 1$ mas	y	5	1932	-47 ± 10	12 ± 9	-10 ± 9	-2 ± 10	-42 ± 9	-18 ± 9
8	$\rho < 1$ mas	n	5	1932	-15 ± 12	2 ± 12	-14 ± 11	-6 ± 12	1 ± 12	11 ± 11
9	$\rho < 2$ mas, $G < 19$	y	5	1382	-57 ± 16	33 ± 13	9 ± 14	3 ± 15	-48 ± 13	-24 ± 14
10	$\rho < 2$ mas, $G < 19$	n	5	1382	-65 ± 20	0 ± 18	22 ± 17	5 ± 20	-30 ± 18	24 ± 17
11a	$\rho < 2$ mas, $[10^5 \alpha] \bmod 2 = 0$	y	5	1255	-19 ± 18	34 ± 15	-10 ± 16	28 ± 17	-10 ± 15	-22 ± 16
11b	$\rho < 2$ mas, $[10^5 \alpha] \bmod 2 = 1$	y	5	1168	-61 ± 17	33 ± 15	17 ± 15	-31 ± 17	-64 ± 15	-18 ± 15

Notes. ρ is the angular separation between the optical and radio positions. N is the number of sources used in the fit. $W = \text{“y”}$ or “n” for weighted or unweighted solution. The weighted solutions use a non-diagonal weight matrix resulting from the combination of *Gaia* covariances and the covariances from the ICRF3-prototype. l_{\max} is the highest degree of the fit from which orientation and glide are extracted for $l = 1$. The columns headed x, y, z give the components of the orientation and glide along the principal axes of the ICRS.

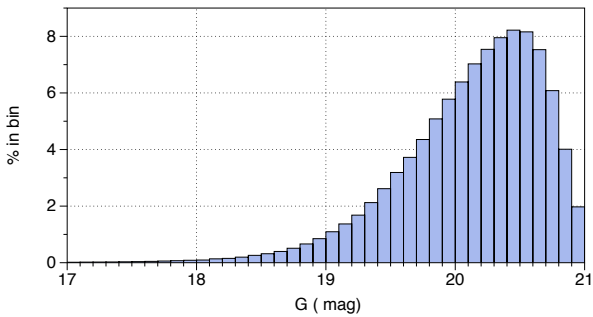


Fig. 20. Magnitude distribution for $\sim 507\,000$ *Gaia* sources that are not included in *Gaia*-CRF2, but are tentatively identified as quasars through a cross-match with the MILLIQUAS catalogue.

Quasars Catalogue (MILLIQUAS; Flesch 2015) is a compilation of quasars and AGNs from the literature, including the release of SDSS-DR14 and AllWISE. We have cross-matched MILLIQUAS² to *Gaia* DR2 using a matching radius of 5 arcsec, but otherwise applying the same selection criteria as for *Gaia*-CRF2. This yielded 1 007 920 sources with good five-parameter solutions in *Gaia* DR2, of which 501 204 are in common with the AllWISE selection in *Gaia*-CRF2. The magnitude distribution of the 506 716 additional sources is shown in Fig. 20. With a median $G \simeq 20.2$ mag, these sources are typically one magnitude fainter than the AllWISE AGNs in *Gaia*-CRF2, with positional uncertainties of about 1 mas.

Obviously, the *Gaia* DR2 release contains even more quasars. They can be found by cross-matching with other QSO catalogues such as the LQAC (Souchay et al. 2015) and various VLBI catalogues. Ultimately, a self-consistent identification of quasars from photometric and astrometric data of *Gaia* will be possible in a future release.

² <http://quasars.org/milliquas.htm>, version of August 2017, containing 1 998 464 entries.

6. Conclusions

With *Gaia* DR2, a long-awaited promise of *Gaia* has come to fruition: the publication of the first full-fledged optical realisation of the ICRS, that is to say, an optical reference frame built only on extragalactic sources. Comprising more than half a million extragalactic sources that are globally positioned on the sky with a median uncertainty of 0.4 mas on average, this represents a major step in the history of non-rotating celestial reference frames built over the centuries by generations of astronomers. The brighter subset with $G < 18$ mag, comprising nearly 30 000 quasars with $\simeq 0.12$ mas astrometric accuracy, is the best reference frame available today and within relatively easy reach for telescopes of moderate size.

We have summarised the detailed content and mapped the main properties of *Gaia*-CRF2 as functions of magnitude and position. The quality claims regarding positional accuracy are supported by independent indicators such as the distribution of parallaxes or proper motions. Large-scale systematics are characterised by means of expansions in vector spherical harmonics. Comparison with VLBI positions in a prototype version of the forthcoming ICRF3 shows a globally satisfactory agreement at the level of 20 to 30 μas . Several sources with significant radio–optical differences of several mas require further investigation on a case-by-case basis.

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