



HAL
open science

Fermi-LAT Observations of High-energy Behind-the-limb Solar Flares

M. Ackermann, A. Allafort, L. Baldini, G. Barbiellini, D. Bastieri, R.
Bellazzini, E. Bissaldi, R. Bonino, E. Bottacini, J. Bregeon, et al.

► **To cite this version:**

M. Ackermann, A. Allafort, L. Baldini, G. Barbiellini, D. Bastieri, et al.. Fermi-LAT Observations of High-energy Behind-the-limb Solar Flares. *The Astrophysical Journal*, 2017, 835 (2), pp.219. 10.3847/1538-4357/835/2/219 . hal-01554386

HAL Id: hal-01554386

<https://hal.science/hal-01554386v1>

Submitted on 3 Jun 2020

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Fermi-LAT Observations of High-energy Behind-the-limb Solar Flares

M. Ackermann, A. Allafort, L. Baldini, G. Barbiellini, D. Bastieri, R.
Bellazzini, E. Bissaldi, R. Bonino, E. Bottacini, J. Bregeon, et al.

► **To cite this version:**

M. Ackermann, A. Allafort, L. Baldini, G. Barbiellini, D. Bastieri, et al.. Fermi-LAT Observations of High-energy Behind-the-limb Solar Flares. *Astrophys.J.*, 2017, 835 (2), pp.219. 10.3847/1538-4357/835/2/219 . hal-01554386

HAL Id: hal-01554386

<https://hal.archives-ouvertes.fr/hal-01554386>

Submitted on 3 Jun 2020

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.



Fermi-LAT Observations of High-energy Behind-the-limb Solar Flares

M. Ackermann¹, A. Allafort², L. Baldini³, G. Barbiellini^{4,5}, D. Bastieri^{6,7}, R. Bellazzini⁸, E. Bissaldi⁹, R. Bonino^{10,11}, E. Bottacini², J. Bregeon¹², P. Bruel¹³, R. Buehler¹, R. A. Cameron², M. Caragiulo^{9,14}, P. A. Caraveo¹⁵, E. Cavazzuti¹⁶, C. Cecchi^{17,18}, E. Charles², S. Ciprini^{16,17}, F. Costanza⁹, S. Cutini^{16,17}, F. D'Ammando^{19,20}, F. de Palma^{9,21}, R. Desiante^{10,22}, S. W. Digel², N. Di Lalla³, M. Di Mauro², L. Di Venere^{9,14}, P. S. Drell², C. Favuzzi^{9,14}, Y. Fukazawa²³, P. Fusco^{9,14}, F. Gargano⁹, N. Giglietto^{9,14}, F. Giordano^{9,14}, M. Giroletti¹⁹, I. A. Grenier²⁴, L. Guillemot^{25,26}, S. Guiriec^{27,50}, T. Jogler²⁸, G. Jóhannesson^{29,30}, L. Kashapova³¹, S. Krucker^{32,33}, M. Kuss⁸, G. La Mura⁷, S. Larsson^{34,35}, L. Latronico¹⁰, J. Li³⁶, W. Liu^{2,37,38}, F. Longo^{4,5}, F. Loparco^{9,14}, P. Lubrano¹⁷, J. D. Magill³⁹, S. Maldera¹⁰, A. Manfreda³, M. N. Mazziotta⁹, W. Mitthumsiri⁴⁰, T. Mizuno⁴¹, M. E. Monzani², A. Morselli⁴², I. V. Moskalenko², M. Negro^{10,11}, E. Nuss¹², T. Ohsugi⁴¹, N. Omodei², E. Orlando², V. Pal'shin⁴³, D. Paneque⁴⁴, J. S. Perkins²⁷, M. Pesce-Rollins⁸, V. Petrosian², F. Piron¹², G. Principe⁴⁵, S. Rainò^{9,14}, R. Rando^{6,7}, M. Razzano^{8,51}, O. Reimer^{2,46}, F. Rubio da Costa², C. Sgrò⁸, D. Simone⁹, E. J. Siskind⁴⁷, F. Spada⁸, G. Spandre⁸, P. Spinelli^{9,14}, H. Tajima^{2,48}, J. B. Thayer², D. F. Torres^{36,49}, E. Troja^{27,39}, and G. Vianello²

¹ Deutsches Elektronen Synchrotron DESY, D-15738 Zeuthen, Germany

² W. W. Hansen Experimental Physics Laboratory, Kavli Institute for Particle Astrophysics and Cosmology, Department of Physics and SLAC National Accelerator Laboratory, Stanford University, Stanford, CA 94305, USA; nicola.omodei@stanford.edu, vahep@stanford.edu

³ Università di Pisa and Istituto Nazionale di Fisica Nucleare, Sezione di Pisa I-56127 Pisa, Italy

⁴ Istituto Nazionale di Fisica Nucleare, Sezione di Trieste, I-34127 Trieste, Italy

⁵ Dipartimento di Fisica, Università di Trieste, I-34127 Trieste, Italy

⁶ Istituto Nazionale di Fisica Nucleare, Sezione di Padova, I-35131 Padova, Italy

⁷ Dipartimento di Fisica e Astronomia "G. Galilei," Università di Padova, I-35131 Padova, Italy

⁸ Istituto Nazionale di Fisica Nucleare, Sezione di Pisa, I-56127 Pisa, Italy; melissa.pesce.rollins@pi.infn.it

⁹ Istituto Nazionale di Fisica Nucleare, Sezione di Bari, I-70126 Bari, Italy

¹⁰ Istituto Nazionale di Fisica Nucleare, Sezione di Torino, I-10125 Torino, Italy

¹¹ Dipartimento di Fisica, Università degli Studi di Torino, I-10125 Torino, Italy

¹² Laboratoire Univers et Particules de Montpellier, Université Montpellier, CNRS/IN2P3, F-34095 Montpellier, France

¹³ Laboratoire Leprince-Ringuet, École polytechnique, CNRS/IN2P3, F-91128 Palaiseau, France

¹⁴ Dipartimento di Fisica "M. Merlin" dell'Università e del Politecnico di Bari, I-70126 Bari, Italy

¹⁵ INFN-Istituto di Astrofisica Spaziale e Fisica Cosmica Milano, via E. Bassini 15, I-20133 Milano, Italy

¹⁶ Agenzia Spaziale Italiana (ASI) Science Data Center, I-00133 Roma, Italy

¹⁷ Istituto Nazionale di Fisica Nucleare, Sezione di Perugia, I-06123 Perugia, Italy

¹⁸ Dipartimento di Fisica, Università degli Studi di Perugia, I-06123 Perugia, Italy

¹⁹ INFN Istituto di Radioastronomia, I-40129 Bologna, Italy

²⁰ Dipartimento di Astronomia, Università di Bologna, I-40127 Bologna, Italy

²¹ Università Telematica Pegaso, Piazza Trieste e Trento, 48, I-80132 Napoli, Italy

²² Università di Udine, I-33100 Udine, Italy

²³ Department of Physical Sciences, Hiroshima University, Higashi-Hiroshima, Hiroshima 739-8526, Japan

²⁴ Laboratoire AIM, CEA-IRFU/CNRS/Université Paris Diderot, Service d'Astrophysique, CEA Saclay, F-91191 Gif sur Yvette, France

²⁵ Laboratoire de Physique et Chimie de l'Environnement et de l'Espace—Université d'Orléans/CNRS, F-45071 Orléans Cedex 02, France

²⁶ Station de radioastronomie de Nançay, Observatoire de Paris, CNRS/INSU, F-18330 Nançay, France

²⁷ NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA

²⁸ Friedrich-Alexander-Universität, Erlangen-Nürnberg, Schlossplatz 4, D-91054 Erlangen, Germany

²⁹ Science Institute, University of Iceland, IS-107 Reykjavik, Iceland

³⁰ NORDITA, Roslagstullsbacken 23, SE-106 91 Stockholm, Sweden

³¹ Institute of Solar-Terrestrial SB RAS, Lermontov st. 126a, 6640333, Irkutsk, Russia

³² University of Applied Sciences and Arts Northwestern Switzerland, CH-5210 Windisch, Switzerland

³³ Space Science Laboratory, University of California, Berkeley, CA 94720-7450, USA

³⁴ Department of Physics, KTH Royal Institute of Technology, AlbaNova, SE-106 91 Stockholm, Sweden

³⁵ The Oskar Klein Centre for Cosmoparticle Physics, AlbaNova, SE-106 91 Stockholm, Sweden

³⁶ Institute of Space Sciences (IEEC-CSIC), Campus UAB, Carrer de Magrans s/n, E-08193 Barcelona, Spain

³⁷ Bay Area Environmental Research Institute, 625 2nd Street, Suite 209, Petaluma, CA 94952, USA

³⁸ Lockheed Martin Solar and Astrophysics Laboratory, 3251 Hanover Street, Bldg. 252, Palo Alto, CA 94304, USA

³⁹ Department of Physics and Department of Astronomy, University of Maryland, College Park, MD 20742, USA

⁴⁰ Department of Physics, Faculty of Science, Mahidol University, Bangkok 10400, Thailand

⁴¹ Hiroshima Astrophysical Science Center, Hiroshima University, Higashi-Hiroshima, Hiroshima 739-8526, Japan

⁴² Istituto Nazionale di Fisica Nucleare, Sezione di Roma "Tor Vergata," I-00133 Roma, Italy

⁴³ St. Peterburg, Vedeneva 2-31, Russia

⁴⁴ Max-Planck-Institut für Physik, D-80805 München, Germany

⁴⁵ Erlangen Centre for Astroparticle Physics, D-91058 Erlangen, Germany

⁴⁶ Institut für Astro- und Teilchenphysik and Institut für Theoretische Physik, Leopold-Franzens-Universität Innsbruck, A-6020 Innsbruck, Austria

⁴⁷ NYCB Real-Time Computing Inc., Lattinatown, NY 11560-1025, USA

⁴⁸ Solar-Terrestrial Environment Laboratory, Nagoya University, Nagoya 464-8601, Japan

⁴⁹ Institutíó Catalana de Recerca i Estudis Avançats (ICREA), E-08010 Barcelona, Spain

Received 2016 October 13; revised 2016 November 23; accepted 2016 December 31; published 2017 January 31

⁵⁰ NASA Postdoctoral Program Fellow, USA.

⁵¹ Funded by contract FIRB-2012-RBFR12PM1F from the Italian Ministry of Education, University and Research (MIUR).

Abstract

We report on the *Fermi*-LAT detection of high-energy emission from the behind-the-limb (BTL) solar flares that occurred on 2013 October 11, and 2014 January 6 and September 1. The *Fermi*-LAT observations are associated with flares from active regions originating behind both the eastern and western limbs, as determined by *STEREO*. All three flares are associated with very fast coronal mass ejections (CMEs) and strong solar energetic particle events. We present updated localizations of the >100 MeV photon emission, hard X-ray (HXR) and EUV images, and broadband spectra from 10 keV to 10 GeV, as well as microwave spectra. We also provide a comparison of the BTL flares detected by *Fermi*-LAT with three on-disk flares and present a study of some of the significant quantities of these flares as an attempt to better understand the acceleration mechanisms at work during these occulted flares. We interpret the HXR emission to be due to electron bremsstrahlung from a *coronal* thin-target loop top with the accelerated electron spectra steepening at semirelativistic energies. The >100 MeV gamma-rays are best described by a pion-decay model resulting from the interaction of protons (and other ions) in a thick-target photospheric source. The protons are believed to have been accelerated (to energies >10 GeV) in the CME environment and precipitate down to the photosphere from the downstream side of the CME shock and landed on the front side of the Sun, away from the original flare site and the HXR emission.

Key words: Sun: flares – Sun: X-rays, gamma rays

1. Introduction

Gamma-ray emission from solar flares is generally believed to occur predominantly in compact high-density regions near the photospheric footpoints of magnetic field lines. Observations of gamma-ray emission from flares whose host active regions (ARs) are located behind the visible solar disk pose interesting questions regarding the acceleration sites and mechanism, the transport and interaction points of the accelerated particles during these rare events.

Three behind-the-limb (BTL) flares with emission up to 100 MeV were observed during solar cycles 21 and 22. The first occulted solar flare whose AR was estimated to be 15° behind the western limb was observed on 1989 September 29 by the Gamma-Ray Spectrometer (GRS) on board the *Solar Maximum Mission* (SMM). Vestrand & Forrest (1993) reported intense gamma-ray line emission in the 1–8 MeV range and a strong 2.223 MeV neutron capture line from this flare. Given the strength of the line emission, it was concluded that a spatially extended component was required in order to explain the observations. The second, detected by PHEBUS on *GRANAT* (Barat et al. 1994) on 1991 June 1, had intense gamma-ray line emission in the 1–8 MeV range but no neutron capture line, indicating that the emission was of coronal origin. The third occulted flare, detected by PHEBUS, BATSE, and EGRET on 1991 June 30, was an electron-dominated flare with no detectable line emission but with significant emission up to almost 100 MeV. Vilmer et al. (1999) report that the spectral properties of this flare were similar to those of the flares occurring on the visible disk. Although there was some speculation and some scenarios were put forth (Cliver et al. 1993), no definite explanations were found on how and where the particles responsible for these emissions were accelerated and where the gamma-rays were produced.

Hard X-ray (HXR) emissions from the loop top of flares originating from ARs located just beyond the limb is often observed by *RHESSI* (Saldanha et al. 2008). These are referred to as partially occulted flares since only the emission from the loop footpoints are occulted. An unusual such flare was reported by Krucker et al. (2007), where the flare originated 40° behind the limb. This is very similar to the 2014 September 1 flare discussed in this paper.

The *Fermi*-Large Area Telescope (LAT) (Atwood et al. 2009) observations have doubled the number of occulted

flares detected and provided the first detections of emission in the GeV range from such rare events. The *Fermi*-LAT observations sample flares from ARs originating from behind both the eastern and western limbs; all flares are associated with very fast coronal massive ejections (CMEs) and strong solar energetic particle (SEP) events. In this paper, we present the observations of the first three BTL flares detected by *Fermi*-LAT together with those from complementary instruments such as *STEREO*, *RHESSI*, *Konus-Wind* (Aptekar et al. 1995), and *Fermi*-GBM (GBM; Meegan et al. 2009), as well as solar observatories in the Radio Solar Telescope Network (RSTN).

In Section 2 we present the multiwavelength features including the light curves (LCs) and localization of the various emissions. Then in Section 3 we describe a more detailed analysis of the spectral data and compare the properties of the BTL flares with three on-disk flares observed by *Fermi*-LAT. In Section 4 we discuss the associated SEP observations, and in Section 5 we present a brief summary and interpretation of these results.

2. Observations of *Fermi*-LAT BTL Flares

We have collected all relevant observations of the three BTL flares detected by *Fermi*-LAT. In what follows we present these observations. A first analysis of the SOL2013-10-11 flare based on Pass7_REP data was presented in Pesce-Rollins et al. (2015). We have reanalyzed all three of the flares using Pass 8 data. The Pass 8 data benefit⁵² from an improved point-spread function, effective area, and energy reach.

2.1. Light Curves

SOL2013-10-11: On 2013 October 11 (Oct13), between 07:05:51 and 07:10:51 UT, *STEREO*-B detected a solar flare with an AR located at N21E103. *GOES* detected an M1.5 class flare starting at around the same time as *STEREO*-B. However, based on the *STEREO*-B 195 Å emission and the method described in Nitta et al. (2013), we estimate that the *GOES* class of this flare had the AR not been occulted would have been M4.9 with an uncertainty within a factor of three. This method utilizes the pre-flare background-subtracted, full-disk integrated EUV intensity, as shown in Figure 1(b). A fast

⁵² A summary of the *Fermi*-LAT Pass 8 performance can be found at http://www.slac.stanford.edu/exp/glast/groups/canda/lat_Performance.htm.

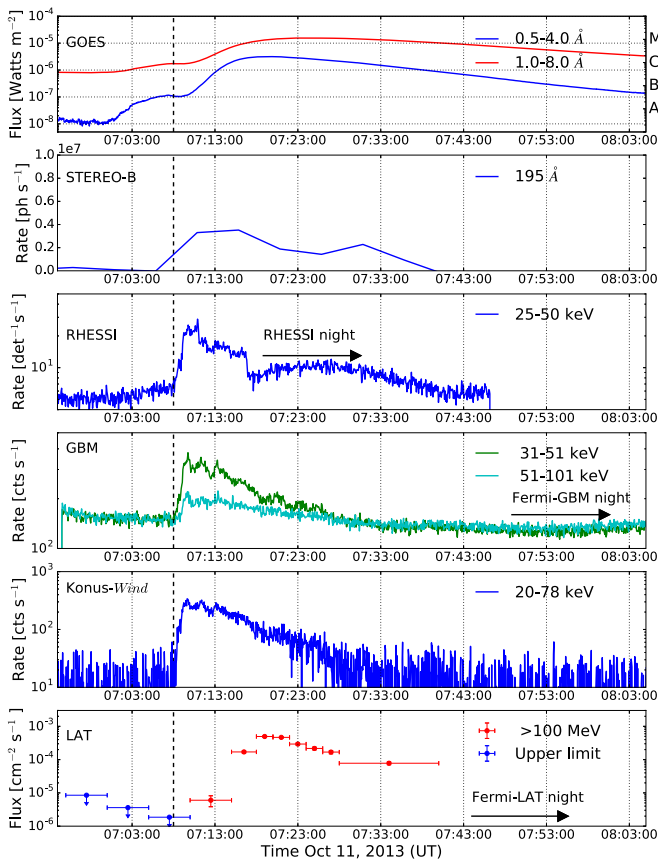


Figure 1. Composite light curve for Oct13 with data from *GOES*, *STEREO*, *RHESSI*, *Fermi*-GBM, *Konus-Wind*, and *Fermi*-LAT. The vertical dashed line represents the estimated start time of the flare, 7:08 UT based on *STEREO*-B observations. *RHESSI* night started at around 07:15 UT and *Fermi*-LAT night started around 07:43 UT.

CME was observed by *LASCO* with a reported first appearance by the white light coronagraph C2 (imaging from 2–6 solar radii) at 07:24:10 UT and a linear speed of 1200 km s^{-1} . *Fermi* coverage started at 07:08:00 UT and continued for more than 30 minutes. *Fermi*-LAT detected $>100 \text{ MeV}$ emission for ~ 30 minute with the maximum of the flux occurring between 07:20:00–07:25:00 UT.⁵³ The *Fermi*-GBM detection of HXRs began a few minutes before *Fermi*-LAT and peaked earlier ($\sim 07:10$ above 50 keV). *RHESSI* coverage was from 07:08:00–07:16:40 UT, overlapping with *Fermi* for 9 minutes. *Konus-Wind*, located at Lagrangian point L1, detected emission in the 20–78 keV and 78–310 keV energy bands simultaneously with *RHESSI* and *Fermi*-GBM. This flare was also observed in radio by the RSTN and the Nobeyama radio polarimeter (Nakajima et al. 1994) at frequencies up to 9 GHz (see Section 3.4).

The LCs for the *GOES*, *STEREO*-B, *RHESSI*, *Fermi*-GBM, *Konus-Wind*, and *Fermi*-LAT are shown in Figure 1. The microwave (MW) emission is compared with the HXR LC from *Konus-Wind* in Figure 2.

SOL2014-01-06: On 2014 January 6 (Jan14), between 07:35:46 and 07:45:46 UT, a solar flare erupted from an AR located at S8W110. Both *STEREO* spacecraft had a full view of the AR and detected a large filament eruption from the AR at

⁵³ From the reanalysis of this flare with Pass 8 data, we gain 5 minutes of additional emission with respect to our previous work (Pesce-Rollins et al. 2015).

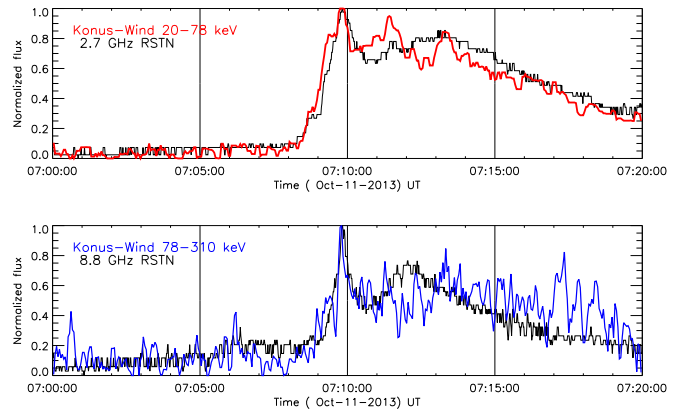


Figure 2. Comparison of the time profiles of HXR and microwave (MW) emission for Oct13. The top and bottom panels compare the 20–78 keV HXR light curve with optically thick gyro-synchrotron (GS) emission at 2.7 GHz and the 78–310 keV HXR light curve with the optically thin GS emission at 8.8 GHz, respectively. Note that the shape of the GS spectrum is such that for frequencies below the peak the spectrum is considered optically thick, whereas above the peak is it optically thin.

approximately 07:50:00 UT. The tip of this filament was seen from the visible solar disk by *SDO*/*AIA*. There is a hint of detection by *GOES* of a smaller-than-C-class flare. However, the peak rate of $2.5 \times 10^5 \text{ photons s}^{-1}$ detected by *STEREO*-B in its 195 Å channel indicates that the flare would have been classified as *GOES* X3.5 had it not been occulted (Nitta et al. 2013). *LASCO* detected a halo CME with a first C2 appearance at 08:00:05 UT with a linear speed of 1400 km s^{-1} .

Upon exiting the South Atlantic Anomaly (SAA)⁵⁴ at 07:55:00 UT, both instruments on board *Fermi* detected emission associated with this flare. *Fermi*-LAT detected $>100 \text{ MeV}$ emission for approximately 20 minutes (with no evidence of temporally extended emission after 08:15:00 UT), and *Fermi*-GBM detected emission in the tens of keV range. *RHESSI* detected emission starting at $\sim 08:18:00$ UT (upon exiting spacecraft night) also in the tens of keV energy range for over 40 minutes. *Konus-Wind* had a full view of the flare and detected emission only in the softest energy band, 20–78 keV, starting at 07:43:00 UT. Radio data above 1.3 GHz from RSTN for this flare did not indicate a detection. The *GOES*, *RHESSI*, *Konus-Wind*, *Fermi*-GBM, and *Fermi*-LAT LCs for this flare are shown in Figure 3.

SOL2014-09-01: On 2014 September 1 (Sep14), between 10:55:56 and 11:00:56 UT, a bright solar flare occurred in an AR located at N14E126. There was no *GOES* signal but *STEREO*-B had an unblocked view of the entire flare and detected a maximum rate of $1.7 \times 10^7 \text{ photons s}^{-1}$ in its 195 Å channel, indicating an unocculted *GOES* X2.4 class flare (Nitta et al. 2013). *LASCO* detected a halo CME with first C2 appearance at 11:12:05 UT with a linear speed of 1900 km s^{-1} . A type II radio burst with an estimated velocity of 2079 km s^{-1} was reported by NOAA Space Weather Alerts in association with this flare. *SDO*/*AIA* reported a coronal wave from this AR starting during 10:45:35–12:21:35 UT. This wave was seen to propagate along the limb and over onto the visible disk.⁵⁵

⁵⁴ Both detectors on board *Fermi* are turned off while the spacecraft is in the SAA.

⁵⁵ As reported by Nitta on <http://www.lmsal.com/isolsearch>.

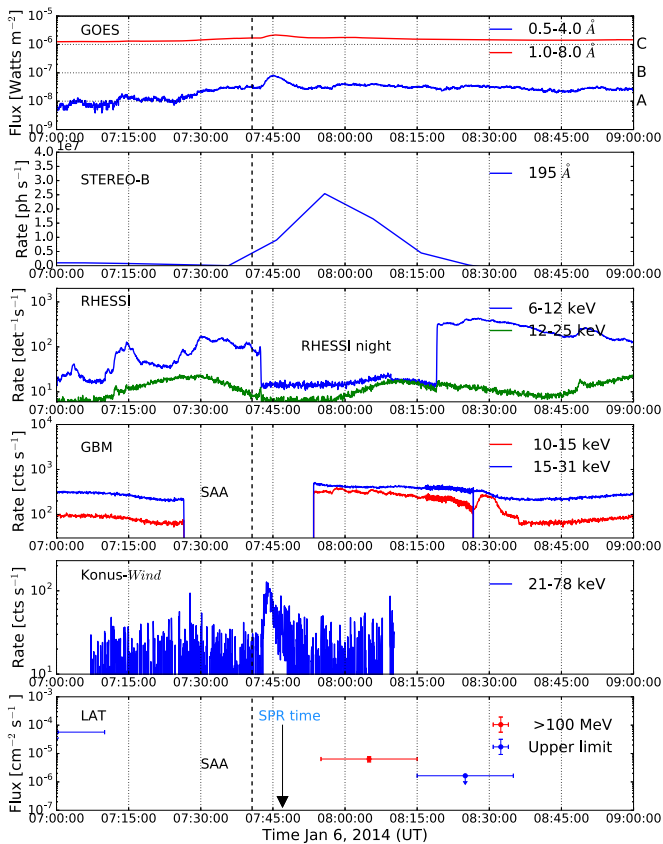


Figure 3. Composite light curve for Jan14 with data from GOES, *STEREO*, *RHESSI*, *Fermi*-GBM, *Konus-Wind*, and *Fermi*-LAT. The vertical dashed line represents the estimated start time of the flare, 07:40:50 UT based on *STEREO*-B emission. The time of Solar Particle Release (SPR) in the solar atmosphere is described in Section 4.

Fermi-LAT detected emission from this flare for ~ 2 hr, peaking between 11:10:00–11:15:00 UT.⁵⁶ The GBM detected emission up to a few MeV in temporal coincidence with the *Fermi*-LAT emission in both the BGO and NaI detectors. *RHESSI* was in the SAA from 10:55:00 to 11:11:00; upon exiting the SAA it detected emission up to 30 keV. *Konus-Wind* detected emission in all three energy bands, 20–78 keV, 78–310 keV, and 310–1180 keV, in temporal coincidence with *Fermi*-GBM. A significant radio flux at frequencies up to 16 GHz was detected by the San Vito station of the RSTN simultaneously with the HXR emission peak detected by both *Fermi*-GBM and *Konus-Wind*. Figure 4 shows the LCs from GOES, *STEREO*, *RHESSI*, *Fermi*-GBM, *Konus-Wind*, and *Fermi*-LAT. The LCs of the microwave emission (MW) are compared with the HXR LC from *Konus-Wind* in Figure 5. The MW intensity distribution with frequency has a maximum at about 1 GHz (see Section 3.4). Thus, we show the time profile of only this frequency for this flare. As evident from Figure 5, there is good agreement between these LCs.

2.2. Localization of the Emission

We present composite images of the Sun as seen by *STEREO* and *SDO* and, whenever available, the position of HXR emission based on *RHESSI* and the >100 MeV gamma-

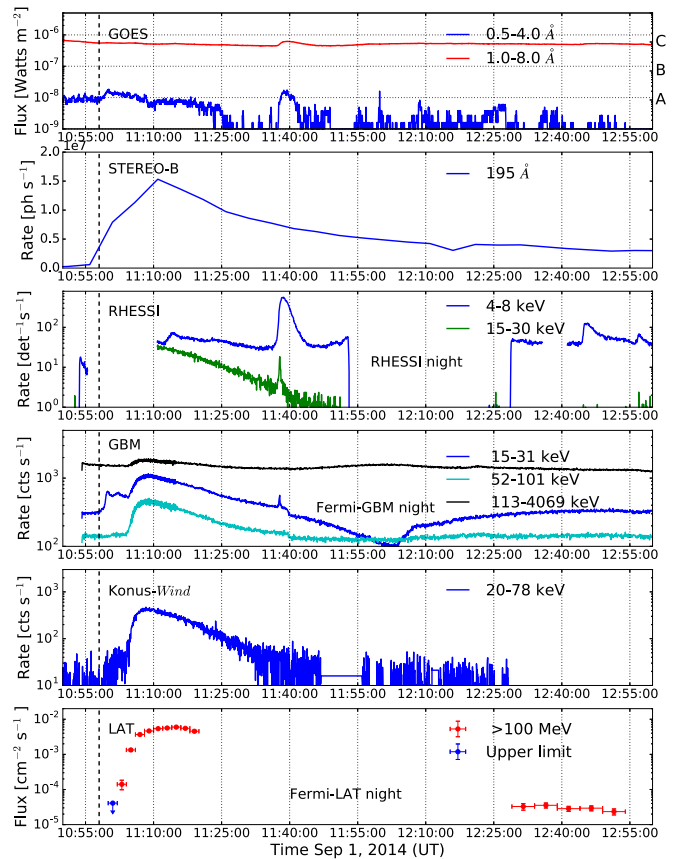


Figure 4. Light curves of Sep14 as seen by GOES, *STEREO*, *RHESSI*, *Fermi*-GBM (black line is for the BGO detector while cyan and blue are for NaI), *Konus-Wind*, and *Fermi*-LAT. The vertical dashed line represents the estimated starting time of the flare, 10:58:00 UT based on *STEREO*-B emission. *Fermi*-LAT night was from $\sim 11:35$ –12:20 UT. The increase in rate seen in *RHESSI* and *Fermi*-GBM at 11:38 UT is due to a small on-disk flare.

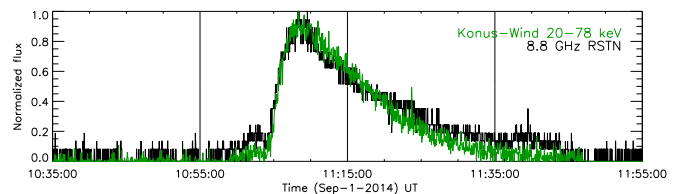


Figure 5. Comparison of the 20–78 keV HXR emission and MW light curves at 8.8 GHz (optically thin gyro-synchrotron emission) for the Sep14 flare. Profiles are normalized to the peak value.

ray emission centroid based on *Fermi*-LAT data. We used the FITS World Coordinate System software package (Thompson & Wei 2010) to co-register the locations of the flares between *STEREO* and *SDO* images. We applied the CLEAN imaging algorithm (Hurford et al. 2002) to *RHESSI* data using detectors 3–9 to reconstruct the X-ray images. The centroid for the >100 MeV gamma-ray emission is determined using the *gtfindsrc* tool, which performs a likelihood analysis of the average position in the time-integrated data set.

Figure 6 displays the composite image for the Oct13 flare showing the *STEREO*-B 195 Å image of the AR located about 10° behind the limb, the *SDO*/AIA 193 Å emission peaking above the limb, a contour of the *RHESSI* image, and results from reanalysis of the flare with Pass 8 data. The new emission centroid is located in heliocentric coordinates $[-880'', 290'']$, about $200''$ closer to the *RHESSI* centroid, and with a 68% error

⁵⁶ Thanks to improvements provided by the new *Fermi*-LAT event selection, we gained 10 extra minutes of coverage with respect to Pass7_REP.

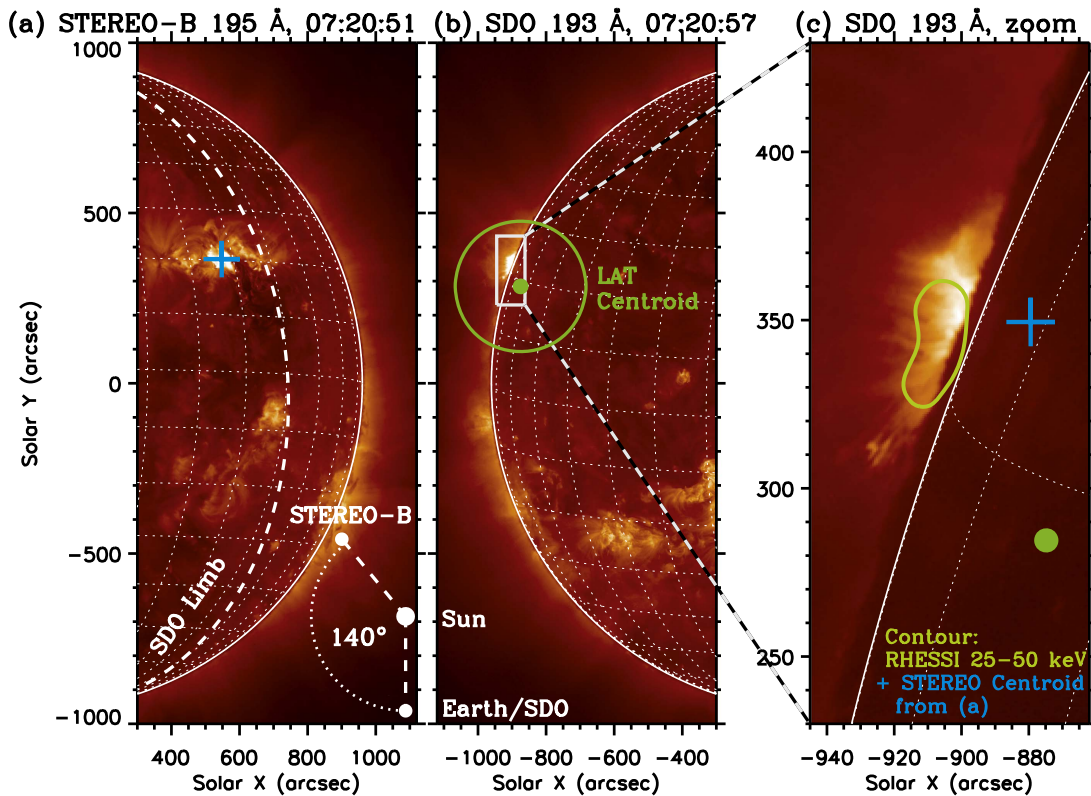


Figure 6. Localization of the Oct13 flare. Images near the flare peak are shown for *STEREO-B* 195 Å (a), *SDO* 193 Å (b), and an enlargement of the *SDO* image (c) marked by the white rectangle in (b). The green circle in (b) shows the 68% error circle for the *Fermi-LAT* emission centroid and the green dot in (c) represents the *Fermi-LAT* emission centroid position. The time range for the *Fermi-LAT* emission is from 07:10:00–07:35:00 UT. The green contour in (c) shows the 25–50 keV *RHESSI* source. The blue cross in (a) and (c) marks the centroid of the *STEREO* flare ribbon as seen from the *STEREO* and Earth/*SDO* perspectives, respectively. The white dashed line in (a) represents the solar limb as seen by *SDO*. The positions of *STEREO* and *SDO*/Earth relative to the Sun are shown in the lower-right corner of (a).

radius of $190''$, which is $\sim 20\%$ smaller than the value we reported in Pesce-Rollins et al. (2015) using Pass7_REP data. In addition, in the Pass 8 data set, the total number of >1 GeV events measured from this flare increased from four to seven. The highest-energy photon detected from this flare was 3.4 GeV and the arrival was 07:19:00 UT.

For the Jan14 flare, the *Fermi-LAT* photon statistics were not sufficient to provide an emission localization error circle smaller than 0.5° . However, we can still conclude that the emission detected by *Fermi-LAT* was consistent with the position of the Sun. In Figure 7 we show the *STEREO-A* and *SDO* images of this event at two different times. The top panels of Figure 7 present *SDO* 171 Å (left) and *STEREO-A* 195 Å (right) image at 07:55:46 UT and show the filament eruption, while the bottom panels show the *SDO* 193 Å (left) and *STEREO-A* 195 Å (right) images at 08:25:46 UT with the *RHESSI* 6–12 and 25–50 keV contours of this flare.

The *Fermi-LAT* >100 MeV emission centroid of Sep14 is located at heliocentric coordinates $[-720'', 610'']$ with a 68% error radius of $100''$. *RHESSI* imaging shows a 6–12 keV source located above the visible limb slightly offset from the *Fermi-LAT* centroid, both shown in Figure 8. If the *RHESSI* source is the loop top of the BTL flare, then the minimum height needed for this source to be visible from $\sim 40^\circ$ BTL would be $\sim 10^{10}$ cm. HXR loop-top emission from a flare located $\sim 40^\circ$ BTL has been detected before by *RHESSI* (Krucker et al. 2007). *Fermi-LAT* measured 17 photons with energies >1 GeV; 15 of these (including a

3.5 GeV photon with an arrival time of 11:16:01 UT) arrived during the first 20 minutes of *Fermi-LAT* detection.

3. Spectral Analysis

3.1. Gamma-Ray Spectra

We performed an unbinned likelihood analysis of the *Fermi-LAT* data with the *gtlike* program distributed with the *Fermi ScienceTools*.⁵⁷ We selected Pass 8 Source class events from a 10° circular region centered on the Sun and within 100° from the local zenith (to reduce contamination from the Earth limb).

We fit three models to the *Fermi-LAT* gamma-ray spectral data. The first two, a pure power law (PL) and a power law with an exponential cutoff (PLEXP), are phenomenological functions that may describe bremsstrahlung emission from relativistic electrons. The third model uses templates based on a detailed study of the gamma-rays produced from the decay of pions originating from accelerated protons with an isotropic pitch angle distribution in a thick-target model (updated from Murphy et al. 1987).

We rely on the likelihood ratio test and the associated test statistic TS (Mattox et al. 1996) to estimate the significance of the detection. The TS is defined as twice the increment of the logarithm of the likelihood obtained by fitting the data with the source and background model components simultaneously.

⁵⁷ We used the version 10-01-00 available from the *Fermi Science Support Center*: <http://fermi.gsfc.nasa.gov/ssc/>.

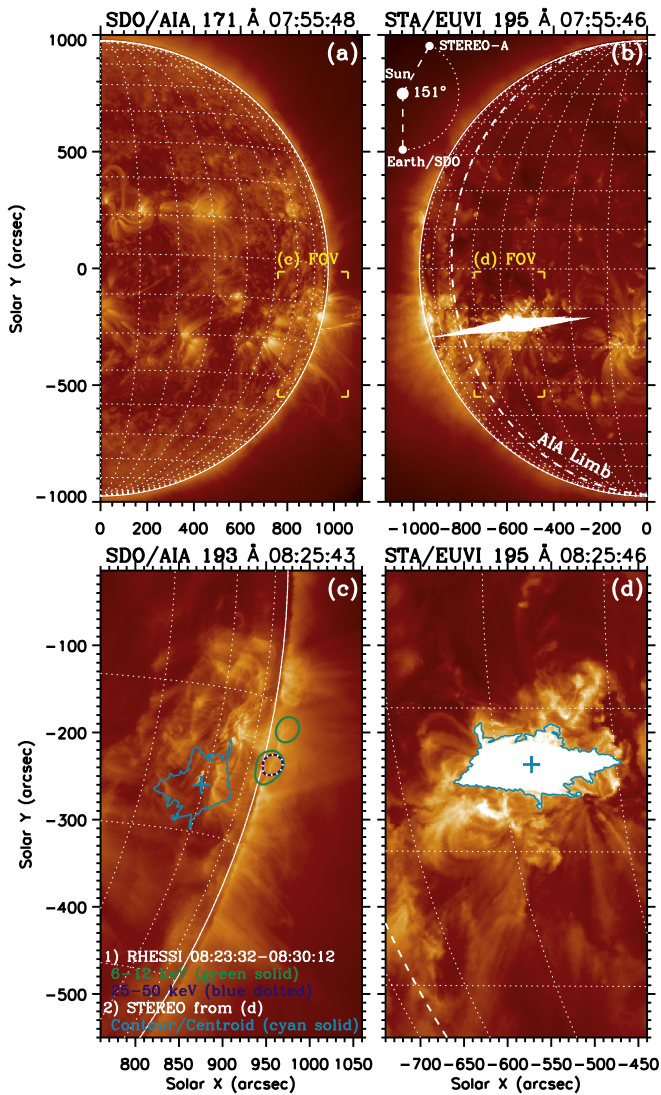


Figure 7. *SDO* 171 and 193 Å (left) and *STEREO-B* 195 Å (right) images near the Jan14 flare peak. The white dashed line in (b) and (d) represents the solar limb as seen by *SDO*. The cyan contour and cross in (d) mark the *STEREO* flare ribbon and its centroid, respectively. Their projected view as seen in the AIA perspective is shown in (c), in which the centroid is located 20° behind the western limb. The green solid and blue dotted contours in (d) show the *RHESSI* 6–12 and 25–50 keV sources, respectively. The rectangular brackets in (a) and (b) mark the field of view for (c) and (d), respectively.

Because the null hypothesis (i.e., the model without an additional source) is the same for the PL and PLEXP models, the increment of the TS ($\Delta\text{TS} = \text{TS}_{\text{PLEXP}} - \text{TS}_{\text{PL}}$) is equivalent to the corresponding difference of the maximum likelihoods computed between the two models. Note that the significance in σ can be roughly approximated as $\sqrt{\text{TS}}$ for two degrees of freedom.

In Table 1 we list the TS_{PL} , ΔTS , photon index Γ for the best-fit model (PL when $\Delta\text{TS} < 25$ or PLEXP when $\Delta\text{TS} \geq 25$), and PLEXP cutoff energy. For several intervals, $\Delta\text{TS} > 25$, indicating that PLEXP provides a significantly better fit than PL. For these intervals we fit the pion-decay models to the data to determine the best proton spectral index following the same procedure described in Ajello et al. (2014). In particular, we performed a series of fits with the pion-decay template models calculated for a range of proton spectral indices. We then fit the resulting profile of the log-likelihood

with a parabolic function of the proton index. The minimum gives the most likely index for the pion-decay model. Note that the TS values for the PLEXP and pion-decay fits cannot be directly compared (Wilks 1938) because they are not nested models. However, PLEXP approximates the shape of the pion-decay spectrum; thus, we expect the pion-decay models to provide a similarly acceptable fit.

The main contributions to the systematic uncertainties are uncertainties in the effective area; in the energy range 100 MeV to ~ 100 GeV, these are of the order of $\pm 5\%$. This uncertainty applies directly to the flux values and from our previous studies on LAT-detected solar flares (Ajello et al. 2014) we find that the systematic uncertainties on the cutoff energy and photon index are also of the order of $\pm 5\%$.⁵⁸

3.1.1. Comparison of BTL and Disk Flare Characteristics

We compare the characteristics of the >100 MeV emission associated with the three BTL flares with three disk flares with similar *GOES* classifications and temporally extended emissions and are described in Ackermann et al. (2014) and Ajello et al. (2014). In addition to spectral parameters, we also compare the total >100 MeV emitted energies, and the total energy released by protons with energy >500 MeV needed to produce the detected gamma-ray emission, based on the templates of (updated from) Murphy et al. (1987). We present these numbers along with the date, estimated *GOES* class, CME speed, AR position, and *Fermi-LAT* detection duration in Table 2. The proton indexes are very similar whereas the on-disk flares appear to have more energy; this is most likely because we observed the on-disk flares over longer timescales. Peak fluxes and the total energy released by protons with $E > 500$ MeV for Sep14 and the impulsive phase of SOL2012-03-07 are comparable.

Figure 9 shows the proton index as a function of time for the Oct13 and Sep14 BTL flares. We fit a constant (red dashed line) and a first-degree polynomial (blue dashed line) to the data. In the case of the constant fit, we list the best-fit value in the upper left corner, whereas for the straight-line fit, we list the value of the slope. The temporal variation over tens of minutes is not sufficient to conclude whether a softening or hardening is present for these BTL flares. In Ajello et al. (2014) we found that for the on-disk flare SOL2012-03-07 the spectrum softened with a timescale of a few hours.

3.2. X-Ray Spectra

As described above we have HXR data from three instruments: *RHESSI*, *Fermi-GBM*, and *Konus-Wind*. *Konus-Wind* was in waiting mode during the Jan14 and Sep14 flares; therefore, these events were detected in only three broad energy channels, namely ~ 20 –78, 78–310, and 310–1180 keV, with 2.944 s time resolution. *RHESSI* provides only a limited coverage of these flares. For the Oct13 flare there is some overlap with *Fermi-GBM* during the rise of the impulsive phase. As shown in Pesce-Rollins et al. (2015) the HXR spectra of *RHESSI* and *Fermi-GBM* agree well. *Fermi-GBM* was in the SAA and *RHESSI* was in spacecraft night during the impulsive phase of the Jan14 flare and detected the flare during the decay phase in only two low-energy channels. Thus, we do

⁵⁸ A more detailed description of the systematic uncertainties tied to the effective area of the LAT can be found here: http://fermi.gsfc.nasa.gov/ssc/data/analysis/LAT_caveats.html.

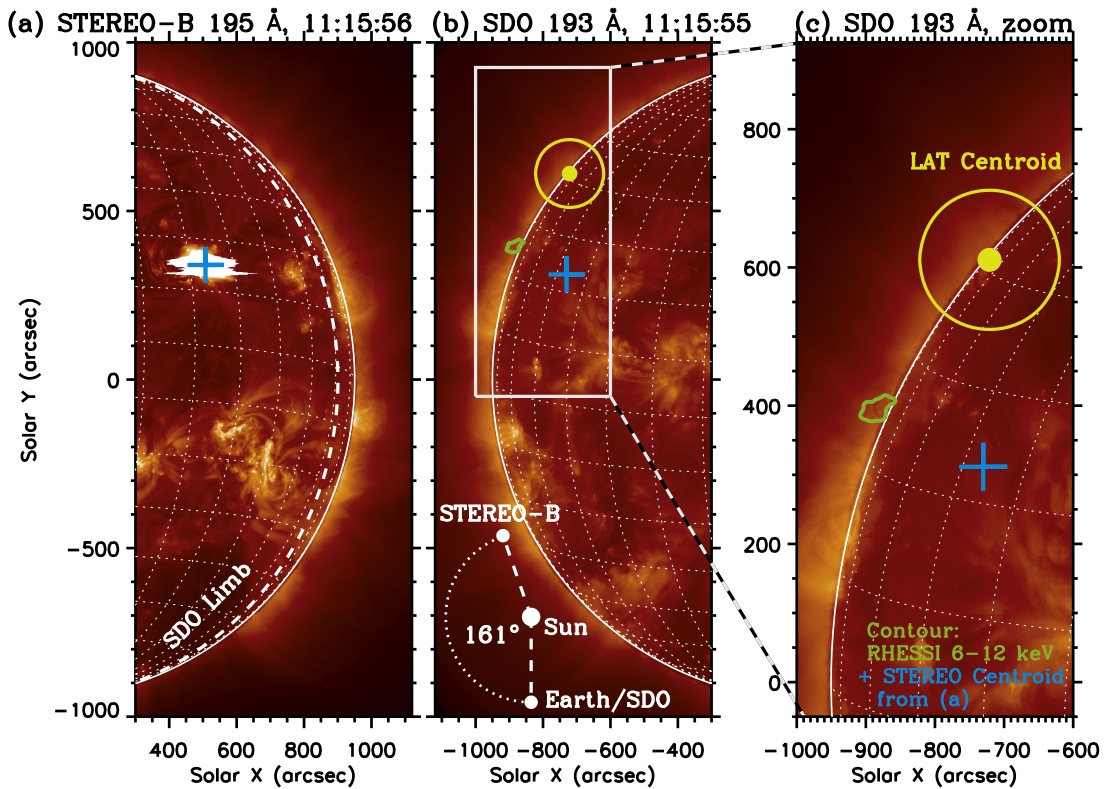


Figure 8. Localization of the Sep14 flare. Images near the time of the flare peak are shown for *STEREO-B* 195 Å (a), *SDO* 193 Å (b), and an enlargement of the *SDO* image (c) marked by the white rectangle in (b). The yellow circle in (b) and (c) shows the 68% error circle for the *Fermi-LAT* emission centroid. The time range of the *Fermi-LAT* data is from 11:02:00–11:20:00 UT. The green contour in (b) shows the 6–12 keV *RHESSI* source. The blue cross in (a) marks the *STEREO* flare ribbon centroid, whose projected position as seen by *SDO* is shown in (b) and (c) with the same symbol. The white dashed line in (a) represents the solar limb as seen by *SDO*.

not have sufficient data for a spectral fit. For the Sep14 flare, *RHESSI* gives information only on the decay phase of the flares. In the following we show results from the analysis of the *Fermi-GBM* data for the Oct13 and Sep14 flares.

3.3. Combined Spectroscopic Studies

For both Oct13 and Sep14, we performed a combined *Fermi-GBM/LAT* fit using the *XSPEC* package (Arnaud 1996). The spectral fits were done by minimizing *PGSTAT*, a profile likelihood statistic that takes into account Poisson error on the total count spectrum and Gaussian error on the background. In order to obtain the background-subtracted spectra of the *Fermi-GBM* data for Sep14, we used both *Fermi-GBM-NaI* and *Fermi-GBM-BGO* spectra accumulated before the flare (from 10:54 UT to 10:57 UT) and after the flare (from 11:42–11:50 UT). For Oct13, because there was a minor on-disk flare whose onset time was 7:01 UT, we used the procedure described in Pesce-Rollins et al. (2015), which consists of using the background estimation tool developed in Fitzpatrick et al. (2012) with an additional 5% systematic error. For the *LAT*, we follow the procedure used in Ajello et al. (2014) and Pesce-Rollins et al. (2015) that consists of deriving the background spectrum directly from the model of the background used in the standard *LAT* likelihood analysis (first using *gtlike* and then *gtbkg*), and obtaining the response of the *LAT* using *gtrspgen* (all of these tools are available in the *FermiScienceTools*). The combined spectral energy distribution (SED) derived from the *Fermi-GBM* and *Fermi-LAT* data extending from 30 keV to 10 GeV for these flares is

shown in Figures 10 and 11, respectively, and the parameters for the spectral fit for both solar flares are listed in Table 3.

The Oct13 flare had a very weak signal in the BGO, and the best-fit model (M_0) consisted of a single PL and the pion-decay templates to describe the bremsstrahlung and >60 MeV emission detected by *Fermi-GBM* and *Fermi-LAT*, respectively. For the Sep14 flare, the best-fit model (M_0) consisted of a single PL with an exponential cutoff at high energies to describe the bremsstrahlung emission, and the pion-decay templates to describe the >60 MeV emission detected by *Fermi-LAT*, similar to what was done for the 2010 June 12 *Fermi* flare (Ackermann et al. 2012). We also tested the statistical significance of an extra PL with a high-energy cutoff (model M_1). To this end, we performed Monte Carlo simulations (using *XSPEC fakeit*) by generating a spectrum described by the best-fit model (M_0). We then re-optimized the parameters of M_0 by fitting it to the simulated data. We finally compared the improvement of *PGSTAT* when fitting the simulated data with the model M_1 . The improvement of *PGSTAT* for the Oct13 and Sep14 flares is $\Delta_{\text{PGSTAT}} \approx 408$ and ≈ 109 , respectively. Despite these high values of Δ_{PGSTAT} , dedicated Monte Carlo simulations indicate that they correspond to significances of 2.5σ and 2.0σ , respectively.

We then checked whether the addition of the 2.223 MeV line (on top of the model M_0) significantly improved the fit. For the Oct13 flare, the improvement was negligible ($\Delta_{\text{PGSTAT}} \approx 0$), while for the Sep14 the improvement was $\Delta_{\text{PGSTAT}} \approx 25$, corresponding to a significance of 2.0σ , estimated from Monte Carlo simulations. We tested whether adding nuclear de-excitation narrow lines and continua provided an improvement

Table 1
Fermi-LAT Spectral Analysis of the Solar Flares Considered in this Work

Time Interval (UT)	TS _{PL}	ΔTS ^a	Photon Index ^b	Cutoff Energy ^c (MeV)	Flux ^d ($\times 10^{-5}$ ph cm ⁻² s ⁻¹)	Proton Index
SOL2013-10-11						
07:10:00–07:15:00	20	3	-2.0 ± 0.4	...	0.9 ± 0.1	...
07:15:00–07:18:00	487	43	-0.4 ± 0.4	130 ± 30	17 ± 1	4.5 ± 0.4
07:18:00–07:20:00	846	77	-0.1 ± 0.4	112 ± 21	49 ± 2	4.3 ± 0.2
07:20:00–07:22:00	776	76	-0.3 ± 0.3	128 ± 23	46 ± 2	4.4 ± 0.3
07:22:00–07:24:00	677	58	-0.2 ± 0.4	153 ± 31	29 ± 2	3.7 ± 0.2
07:24:00–07:26:00	345	25	-0.9 ± 0.4	208 ± 61	21 ± 2	3.9 ± 0.3
07:26:00–07:28:00	219	34	0.6 ± 0.7	86 ± 22	16 ± 1	4.5 ± 0.4
07:28:00–07:40:00	283	29	0.1 ± 0.7	86 ± 26	7 ± 1	5.3 ± 0.4
SOL2014-01-06						
07:55–08:15	67	20	-2.4 ± 0.2	...	0.6 ± 0.1	...
SOL2014-09-01						
11:02:00–11:04:00	42	9	-2.5 ± 0.3	...	11 ± 3	...
11:04:00–11:06:00	321	41	-0.2 ± 0.5	117 ± 27	117 ± 10	4.7 ± 0.4
11:06:00–11:08:00	1070	120	-0.5 ± 0.3	116 ± 15	360 ± 17	5.2 ± 0.2
11:08:00–11:10:00	1549	126	-0.9 ± 0.2	158 ± 20	477 ± 19	4.9 ± 0.2
11:10:00–11:12:00	2549	115	-1.2 ± 0.2	194 ± 28	565 ± 21	4.8 ± 0.2
11:12:00–11:14:00	6394	157	-0.8 ± 0.1	167 ± 10	522 ± 20	4.4 ± 0.1
11:14:00–11:16:00	2430	159	-0.6 ± 0.2	148 ± 17	540 ± 22	4.5 ± 0.2
11:16:00–11:18:00	1069	32	-1.9 ± 0.2	488 ± 90	465 ± 23	4.2 ± 0.2
11:18:00–11:20:00	1047	49	-1.0 ± 0.3	182 ± 37	396 ± 27	4.5 ± 0.2
12:29:00–12:34:00	56	12	-2.2 ± 0.2	...	3 ± 1	...
12:34:00–12:39:00	133	3	-2.3 ± 0.2	...	3 ± 1	...
12:39:00–12:44:00	130	5	-2.2 ± 0.2	...	2 ± 1	...
12:44:00–12:49:00	97	11	-2.3 ± 0.2	...	2 ± 1	...
12:49:00–12:54:00	84	3	-2.4 ± 0.2	...	2 ± 1	...

Notes.

^a $\Delta TS = TS_{\text{PLEXP}} - TS_{\text{PL}}$.

^b Photon index from the best-fit model. The PL is defined as $\frac{dN(E)}{dE} = N_0 E^\Gamma$ and the PLEXP as $\frac{dN(E)}{dE} = N_0 E^\Gamma \exp(-\frac{E}{E_c})$, where E_c is the cutoff energy.

^c From the fit with the PLEXP model.

^d Integrated flux between 100 MeV and 10 GeV calculated for the best-fit model.

to the fit for both flares and found that it was not significant. The nuclear line and continua templates used here are based on a detailed study of the nuclear gamma-ray production from accelerated ion interactions with elements found in the solar atmosphere (Murphy et al. 2009).

3.4. Radio Spectra

The radio spectra obtained at the peaks of the Oct13 and Sep14 events are shown in Figure 12. The spectrum was obtained by using the RSTN one-second resolution fluxes. The spectra were background subtracted and the fluxes were integrated within the selected time intervals for each frequency. For each spectrum we detected the frequency where the maximum flux was observed. The spectra of the Oct13 and Sep14 flares are similar to those expected from the gyro-synchrotron (GS) mechanism, with peaks that separate the optically thin and thick parts (for GS spectrum properties see, for example, Dulk 1985). The values for the microwave photon spectral index, α , for both flares are shown in Figure 12. Unfortunately, the RSTN data for the time period of the Jan14 flare are not sufficient to obtain radio spectra.

4. Solar Energetic Particles

Interpretation of the origin of SEPs is more complicated than studying photons, since they move with different velocities along curved magnetic field lines and are detected only when the field lines are connected to the instrument. Moreover, they may be scattered by turbulence with energy-dependent mean free path. The *GOES* proton fluxes did not show a significant increase for the Oct13 event (which occurred behind the eastern limb). However, *STEREO-B*, positioned behind the limb and with better magnetic connectivity, detected an increase of SEP proton intensity starting roughly 2 hr after the detection of an electromagnetic signal from this flare. The Sep14 flare, however, which was located roughly 40° behind the eastern limb, was associated with an increase in the proton fluxes detected both by *STEREO-B* (which had a front view of the flare) and *GOES* (starting roughly 9 hr after the flare).

The Jan14 flare, originating in an AR 20° behind the western limb, produced no significant increase in the proton flux at *STEREO*, had better magnetic connectivity to the Earth, and was associated with a very strong SEP event with neutrons detected on the ground by the South Pole neutron monitors.

The arrival time T_2 (Debrunner et al. 1988) at the Earth of a particle with velocity $v = \beta c$ released from the Sun at time T_1 ,

Table 2
Comparison Between Behind-the-limb and On-disk Flare Quantities

Date (UTC)	GOES ^a Class	CME Speed ^b (km s ⁻¹)	AR Position	Duration (minutes)	Peak Flux ^c (10 ⁻⁵ ph cm ⁻² s ⁻¹)	$E_\gamma > 100$ MeV ^d (erg)	Proton Index ^e	$E_p > 500$ MeV ^f (erg)
2013-10-11	M4.9	1200	N21E103	30	49 ± 2	1.5 × 10 ²³	4.3 ± 0.1	9.8 × 10 ²⁴
2014-01-06	X3.5	1400	S8W110	20	0.8 ± 0.1	4.2 × 10 ²¹	5.3 ± 0.4 ^g	3.5 × 10 ²³
2014-09-01	X2.4	2000	N14E126	113	565 ± 14	1.4 × 10 ²⁴	4.7 ± 0.1	7.0 × 10 ²⁵
On-disk flares								
2011-03-07	M3.7	2125	N30W48	798	3 ± 1	5.1 × 10 ²³	4.7 ± 0.2	3.6 × 10 ²⁵
2011-06-07	M2.5	1255	S22W53	38	3 ± 1	3.2 × 10 ²²	5.0 ± 0.3	2.5 × 10 ²⁴
2012-03-07 ^h	X5.4	2684	N16E30	45	417 ± 13	3.9 × 10 ²⁴	3.90 ± 0.02	2.1 × 10 ²⁶
2012-03-07E ⁱ	X5.4	2684	N16E30	1068	97 ± 2	1.4 × 10 ²⁵	4.3 ± 0.1	9.0 × 10 ²⁶

Notes.

^a GOES class for BTL flares is estimated based on the *STEREO* 195 Å flux.

^b Speed is the linear speed reported by the *LASCO* CME catalog.

^c For photon energies >100 MeV.

^d Total energy released in >100 MeV gamma-rays integrated over the time interval when $\Delta T_S > 25$.

^e Proton index in the same time interval as (d).

^f Total energy released by protons with $E > 500$ MeV estimated over the same interval as in (d). Values may be underestimated for flares with centroids at heliocentric angles $>75^\circ$.

^g The ΔT_S value for this flare is 20, so the improvement over a power law is marginal.

^h Impulsive phase of the flare, *Fermi*-LAT detection from 00:38:52-01:23:52 UT. Note that the peak of the *GOES* X-ray flare occurred ~ 6 minutes prior to *Fermi* orbital sunrise.

ⁱ Extended phase of the flare, starting from 02:27 UT.

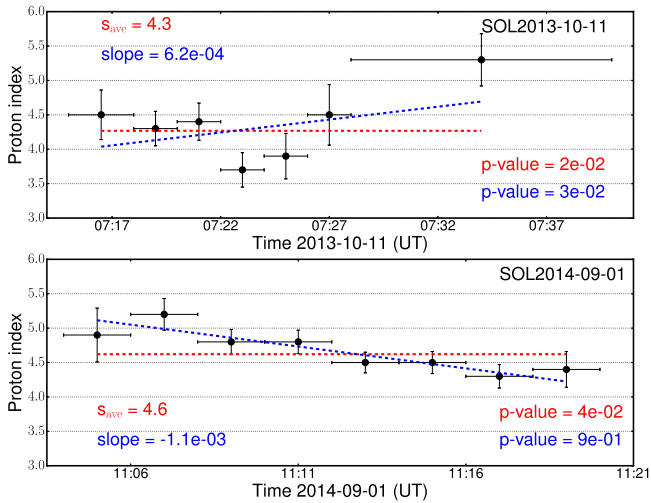


Figure 9. Proton index inferred from the fitting of the gamma-ray emission with the pion-decay templates as a function of time for Oct13 (top panel) and Sep14 (bottom panel). The red dashed lines represent a fit with a constant and the blue dashed lines a linear fit. Fit parameter values and p -values are indicated in each panel.

is

$$T_2 \left(\frac{1}{\beta} \right) = T_1 + \frac{\zeta}{c} \times \frac{1}{\beta}, \quad (1)$$

where ζ is the distance traveled by the SEP from the acceleration region to the Earth. Assuming scatter-free propagation for the first-arriving particles, ζ will be independent of particle energy or velocity and will be equal to the path length along a Parker spiral magnetic field line from the Sun to the Earth (usually assumed to be 1.2 au). If all the particles are accelerated at the same place and time, then the intercept of the line fitting T_2 to $1/\beta$ would give T_1 , the time at which SEPs are

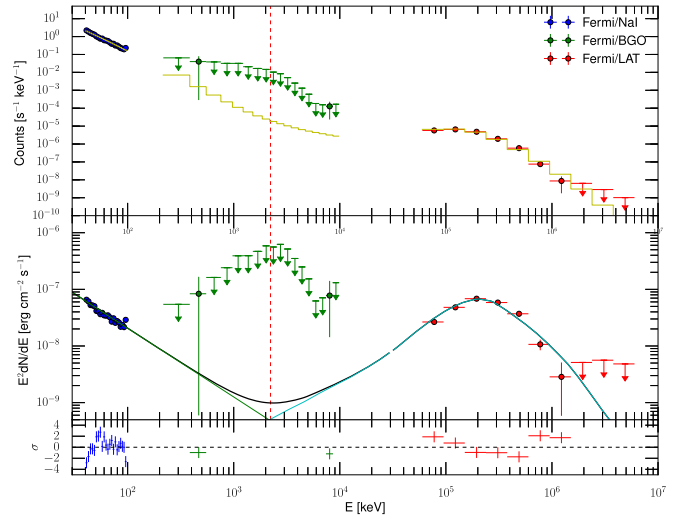


Figure 10. Combined *Fermi*-GBM/LAT count spectra (top panel) and spectral energy distribution (middle panel) for the Oct13 flare integrated between 7:10 and 7:35 UT; the lower panel shows the residuals of the fit. The models that best fit the data are a power law that dominates the low-energy spectrum and a pion-decay model, which describes the *Fermi*-LAT spectrum. The neutron capture line (at 2.223 MeV, highlighted by the red vertical dashed line) is not statistically significant, and neither is an additional low-energy power law with exponential cutoff ($\sim 2.5\sigma$). The parameters of the fits are listed in Table 3.

released and is known as the Solar Particle Release (SPR) time. The slope of this line, ζ/c , will be of the order of 600 s.

Since we do not have high energy resolution SEP data we obtain the SEP onset times using two methods depending on the energy range of the particles. For the GOES SEPs with energies <200 MeV we apply a median filter to the intensity profile and define the onset time as when the profile reached 5% of the maximum. For the higher-energy SEPs we apply a Fast Fourier Transform (FFT) filter to the intensity profile and evaluate the time when the second derivative is maximum as

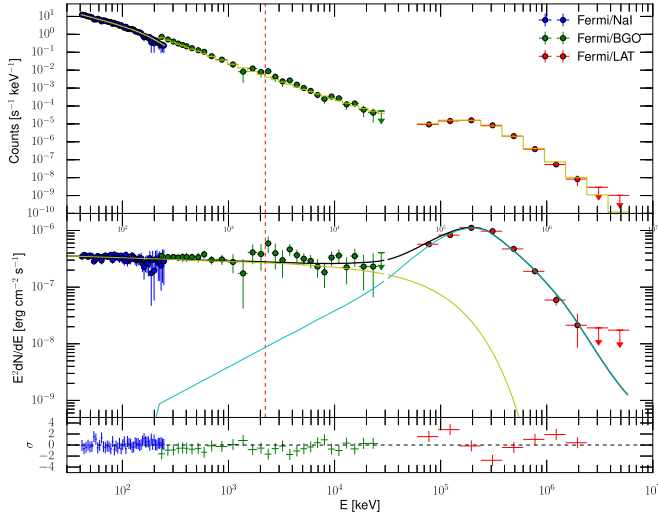


Figure 11. Combined *Fermi*-GBM/LAT count spectra (top panel) and SED (middle panel) for the Sep14 flare integrated between 11:02 and 11:20 UT; the lower panel shows the residuals of the fit. The best-fit models are a power law with exponential decay at high energy to describe the emission from 30 keV to ~ 10 MeV and a pion-decay model to describe the *Fermi*-LAT spectrum. The neutron capture line (at 2.223 MeV, highlighted by the red vertical dashed line) is not statistically significant ($\sim 2\sigma$) and neither is an additional power law at low energy $\sim 2\sigma$. The fit parameters are listed in Table 3.

Table 3

Best-fitting Spectral Parameters of the *Fermi*-LAT and *Fermi*-GBM Data

Parameter	Value
SOL2013-10-11	
Best-fit model: PL1+PION	
Power-law (1) index	3.2 ± 0.1
Proton Index	4.1 ± 0.1
PL1+PL2×EXP+PION	
Power-law (1) index	3.4 ± 0.06
Power-law (2) index	-0.9 ± 0.3
Cutoff energy (MeV)	0.8 ± 0.1
Proton Index	4.1 ± 0.1
SOL2014-09-01	
Best-fit model: PL1×EXP+PION	
Power-law (1) index	2.06 ± 0.01
Cutoff energy (MeV)	90 ± 7
Proton Index	4.4 ± 0.1
(PL1+PL2)×EXP+PION	
Power-law (1) index	2.18 ± 0.01
Power-law (2) index	1.4 ± 0.3
Cutoff energy (MeV)	10 ± 0.1
Proton Index	4.4 ± 0.1

Note. Model parameters for both the Oct13 and Sep14 flares. For the Oct13 flare we integrated between 7:10 and 7:35 UT, and for the Sep14 flare we integrated from 11:02 and 11:20 UT. The parameters of the best-fit model are compared with those of a more complex model with double power laws at low energy (whose improvement is not statistically significant).

the onset time. To estimate of the uncertainty on the onset time we performed a scan over a series of values for the median and FFT filter windows and take the difference in onset times as the error.

The onset times as a function of $\frac{1}{\beta}$ for the *GOES* SEP protons associated with Jan14 are shown in Figure 13. Fitting the data

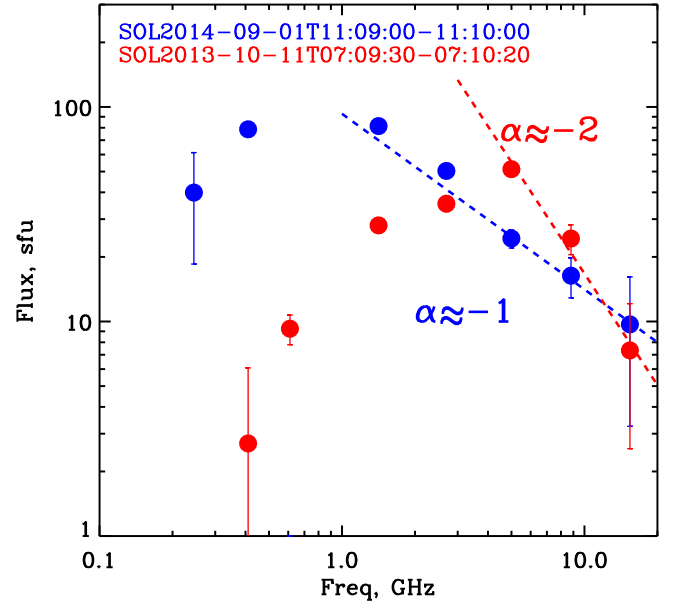


Figure 12. Radio spectra during the peaks of the HXR and radio time profiles (07:09:30–07:10:20 UT) for the Oct13 flare in red and for the Sep14 flare in blue (11:09:00–11:10:00 UT). The microwave photon spectral index α , found by fitting the optically thin part of the spectra for each flare, is also shown.

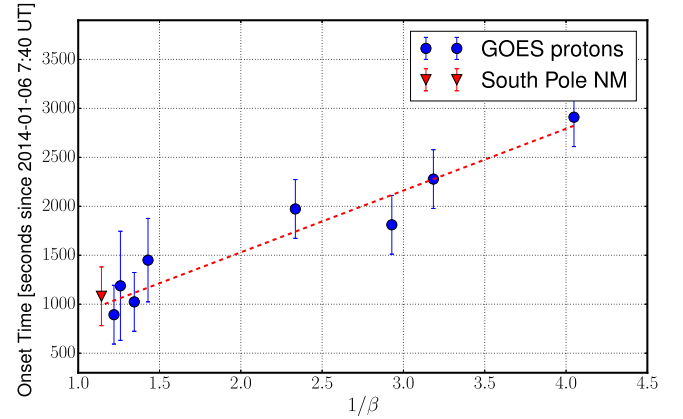


Figure 13. Onset times as a function of $1/\beta$ for the *GOES* SEP protons associated with Jan14. The Y-axis gives the time in seconds since 2014-01-06 07:40 UT (assumed start of the flare as detected by *STEREO*-B) assuming scatter-free propagation for the first-arriving particles. The red dashed line represents the fit using Equation (1), giving the path length $\zeta = 1.3 \pm 0.2$ au and $T_1 = 07:55 \pm 00:05$ UT.

to Equation (1), we estimate the SPR time to be $T_1 = 07:55 \pm 0:05$ UT and the path length $\zeta \sim 1.3 \pm 0.2$ au.

The value for T_1 is consistent with the estimated SPR time, 07:47 UT, reported by Thakur et al. (2014) but is somewhat later than the start of the flare activity at the Sun. Since *Fermi* was in the SAA and *RHESSI* was in orbital night, the most accurate time of the initiation of the flare was obtained from *Konus-Wind* and *STEREO* data. *Konus-Wind* detected emission in the 21–78 keV energy range starting at 07:43 UT, implying an emission time at the Sun of 07:35 UT, roughly 12 minutes prior to the SPR time. Assuming that the CME was released at this latter time, the acceleration of SEPs started when the CME was at a height of $1.5 R_\odot$ ($v_{\text{CME}}/1400 \text{ km s}^{-1}$).

Unfortunately, the high-energy SEP data were not sufficient to estimate the SPR time for the other two BTL flares that occurred on Oct13 and Sep14.

5. Summary and Discussion

We have presented the analysis of *Fermi*-LAT data from three BTL solar flares, whose location is determined by *STEREO* observations. We also presented analyses of data from *RHESSI*, *Fermi*-GBM, and *Konus-Wind* in the HXR range; *SDO* in EUV; RSTN in radio wavelengths; some SEP observations; *GOES*; and the South Pole Neutron Monitors. We found several results.

1. The HXR LCs measured by the three instruments (in overlapping energy ranges) are in good agreement with each other and with the high-frequency radio LCs. The *Fermi*-LAT emission, in general, commences several minutes after the HXRs, peaks at a later time, and lasts longer.
2. We are able to obtain *RHESSI* images in soft HXRs for flares just over the limb, indicating a loop-top emission from loops with heights above the photosphere ranging from 0.2×10^{10} to 2×10^{10} cm. *SDO* detected UV emission for Oct13, and possibly Jan14, but not for the Sep14 flare. The Jan14 flare was too weak to permit localization of the gamma-ray source, but localizations of the *Fermi*-LAT emission with Pass 8 data for the other two flares situates the centroid locations near the solar limb and about $50''$ (Oct13) and $300''$ (Sep14) from the *RHESSI* centroid.
3. The LCs (available only for the Oct13 and Sep14 flares) show good agreement between the HXR spectra measured by *Fermi*-GBM, *RHESSI* (whenever contemporaneous data are available), and *Konus-Wind* below a few 100 keV. The *Fermi*-LAT spectra can be fitted with a relatively hard PLEXP phenomenological model (possibly due to relativistic electron bremsstrahlung) or to a pion-decay model with proton indexes around 4. We note that production of photons of about ~ 3 GeV detected here requires >15 GeV protons or >6 GeV electrons. These energies are greater than any other reported for solar particles.
4. We compare several spectral properties of >100 MeV radiation from the BTL flares with those of three on-disk flares previously analyzed. As evident from Table 2 and Figure 9, in general, these flares have similar characteristics. Similar *GOES* class flares have similar peak photon fluxes, surprisingly, regardless of limb occultation. The higher total gamma-ray energy of on-disk flares most likely results from their having been observed over longer timescales. Thus, the underlying acceleration and emission processes are most likely similar, but the transport paths of the radiating particles (presumably protons) from the acceleration to the emission site must be different. That is, for BTL flares, protons must travel greater distances to land on the front side of the Sun to produce the detected gamma-rays.
5. We find that the proton index does not vary significantly during the 30 minute observation time windows for the BTL flares. This is similar to that found for the on-disk flares for the same time frames, but some on-disk flares are detected over longer time intervals (up to 20 hr) during which the proton index increases gradually.

We now present our interpretation of these results with the goal of constraining the emission and acceleration mechanisms and sites.

5.1. SEP Onset Times

Since the spectral information available for the Jan14 flare is limited, we cannot constrain the acceleration, transport, or radiation processes. However, this flare is remarkable because the AR, being located beyond the western limb, provided good magnetic connection with the Earth and as a result it was associated with a very strong SEP event and increased the neutron flux in the South Pole neutron monitor. As presented in Figure 13, the acceleration time at the Sun for the SEP protons is $07:55 \pm 0:05$ UT, which is in agreement with the SPR time of $07:47 \pm 0:08$ UT reported by Thakur et al. (2014). Unfortunately, *Fermi*-LAT was in the SAA at the SPR time and exited at $07:55$ UT. Therefore, we can only provide an upper limit on the time of acceleration of gamma-ray-producing particles (presumably protons) at the Sun. However, *Konus-Wind* detected emission in the 20–78 keV band starting at $07:43$ UT, implying that the electrons were accelerated at the Sun at $07:35$ UT, roughly 12 minutes prior to the SPR.

Due to the location of the AR, the other two flares were not magnetically well connected to the Earth. For the Sep14 flare, the *GOES* proton fluxes begin to increase roughly 9 hr after the flare start time. *STEREO*-B observed the flare in its entirety but provides proton fluxes only in three energy bands, which is not sufficient to estimate the acceleration time at the Sun as we did for Jan14. Thus, we do not have reliable SEP data to test whether the onset times of the SEP and gamma-rays coincide; however, we can compare the height of the CME at the time of the gamma-ray onset with the CME height at SPR found for Jan14. Based on the second-order fit of height as a function of time (from the LASCO CME online catalog), we estimated that the CME was at $\sim 2.5 R_{\odot}$ at the onset time of the *Fermi*-LAT emission. We found a similar value for Oct13. These heights are significantly smaller than the predicted values reported in Thakur et al. (2014) for the CME heights at SPR as a function of source longitudes from Solar Cycle 23. To test whether these CME heights are compatible with the gamma-ray emission detected from the BTL flares, detailed simulations on particle transport in the solar magnetic field are required. This study is beyond the scope of this paper and will be addressed in a future work.

5.2. Spatial Distributions

The *RHESSI* images show that the 20–50 keV HXRs, due to electron bremsstrahlung, are emitted from $\approx 30''$ ($L \sim 2 \times 10^9$ cm) and $\approx 50''$ ($L \sim 3.5 \times 10^9$ cm) sources from the tops of loops extending to heights of $\sim 4 \times 10^9$ and $\sim 2 \times 10^{10}$ cm above the photosphere for the Oct13 and Sep14 flares, respectively. The column depth required for a particle of mass m and kinetic energy E (in units of mc^2), or Lorentz factor $\gamma = E + 1 = (1 - \beta^2)^{-1/2}$, to lose all its energy via Coulomb interactions is

$$\begin{aligned} N_{\text{Coul}} &= (4\pi r_0^2 \ln \Lambda)^{-1} (m/m_e) \int_0^E \beta^2 dE \\ &= 5 \times 10^{22} (m/m_e) E^2 / \gamma \text{ cm}^{-2}, \end{aligned} \quad (2)$$

where r_0 is the classical electron radius and $\ln \Lambda \sim 20$ is the Coulomb logarithm. From Equation (2) we find that the column depth required to stop the >50 keV electrons that produce the HXRs detected by *RHESSI* is $\sim 5 \times 10^{20} \text{ cm}^{-2}$. Assuming a density of $n < 10^{10} \text{ cm}^{-3}$, the sizes of the loop-top sources imply a column depth of $N_{\text{LT}} < 2$ and $3.5 \times 10^{19} \text{ cm}^{-2}$ for the

two HXR sources, respectively. This means that the energy-loss time $\tau_{\text{Coul}} \sim 2N_{\text{Coul}}/(vn)$ is more than 10 times longer than the crossing time $\tau_{\text{cross}} \equiv L/v = N_{\text{LT}}/(vn)$. These values, together with the results reported in Chen & Petrosian (2013), suggest that the HXR emission from these two BTL flares are thin-target sources.

If these loop-top sources were to be produced in a thick target, which can come about in two ways, it would be under extreme conditions not likely to be the case here. The first is in the strong diffusion limit, i.e., if the electrons are trapped because the scattering time, τ_{sc} , is smaller than the crossing time, τ_{cross} . Note that because $\tau_{\text{Coul}} > \tau_{\text{cross}}$, the scattering cannot be due to Coulomb interactions but it could be due to turbulence. The second possibility is in the weak diffusion limit $\tau_{\text{sc}} > \tau_{\text{cross}}$, where trapping can occur if the field lines converge strongly. In this case we need $\tau_{\text{sc}} > \tau_{\text{Coul}}/\eta$ ⁵⁹ (with proportionality constant η increasing with increasing field convergence (Petrosian 2016)). These are somewhat extreme conditions requiring either a very high density (therefore increasing the scattering, presumably turbulence) or a somewhat strong field convergence and a “clean” loop with low density and low level of turbulence.

Thus, the HXR emission is most likely due to electron bremsstrahlung from a thin-target loop-top source.

Based on the positions of the *Fermi*-LAT >100 MeV emission centroids alone, we cannot exclude that the gamma-rays come from the loop-top source. In order to investigate this possibility, we rely on Equation (2) to find the column depth required to stop >350 MeV protons (which produce >100 MeV photons) and find $N_{\text{Coul}}(350 \text{ MeV}) \sim 10^{25} \text{ cm}^{-2}$. This is much larger than the loop-top column depth $N_{\text{LT}} \sim 2\text{--}3.5 \times 10^{19} \text{ cm}^{-2}$ so that, in the absence of trapping, we are again dealing with a thin-target loop-top source and a small energy loss compared to what would be the case for emission from a thick-target photospheric source. Consequently, we would need much larger energies of accelerated protons than those given in Table 2 assuming a thick-target scenario. The condition for proton trapping in the loop top (or any coronal trap region, for that matter) is also more extreme than that described above for electrons, and we would need escape times 10^5 times longer than the crossing time. Thus, the emission detected by the *Fermi*-LAT is most likely due to decay of pions produced by energetic protons interacting in a thick-target photospheric source.

5.3. Spectral Distributions

As shown in Lin & Hudson (1971) and McTiernan & Petrosian (1990), in a thin-target scenario the relation between the HXR spectral index, γ , and the number index of the energy spectrum of the emitting electrons⁶⁰ through bremsstrahlung processes, δ , is

$$\delta = \begin{cases} \gamma - 0.5 & \text{non-relativistic case,} \\ \gamma & \text{relativistic case.} \end{cases} \quad (3)$$

⁵⁹ According to Mal'ushkin & Kulsrud (2001), $\eta \sim \ln(B_0/B_L)$, the log of the ratio of magnetic field strengths at the middle and ends of a magnetic bottle.

⁶⁰ This should not be confused with the electron flux, $F(E) = vN(E)$, index, which in the nonrelativistic limit is $\delta \ln F(E)/d \ln E = \gamma - 1.0$, used commonly. In the relativistic regime these indexes are equal and as stated in Equation (3) $\delta = \gamma$.

For optically thin synchrotron radiation, the relation between δ and the microwave photon spectral index, α , can be described by the following relations for the relativistic (e.g., Rybicki & Lightman 1979) and semirelativistic (e.g., Dulk 1985) cases

$$\delta = \begin{cases} 1.1\alpha + 1.36 & \text{semi-relativistic case,} \\ 2.0\alpha + 1 & \text{relativistic case.} \end{cases} \quad (4)$$

We will use these relations in the following for the interpretation of the detected emission from the two brightest BTL flares reported in this work. The spectra of the two flares in the 30 keV–10 MeV energy range are somewhat different so we discuss them separately.

Oct13: As already shown in Pesce-Rollins et al. (2015), both *RHESSI* and *Fermi*-GBM detect HXR spectra for the Oct13 flare that are relatively steep (the best-fit value of the power-law photon index $\gamma = 3.2 \pm 0.1$). However, as shown in Figure 10, there is a tentative indication that the spectrum may harden into a nearly flat part in the 1–10 MeV range, requiring either a similar hardening of the spectrum of the accelerated electrons, as, e.g., described in Park et al. (1997), or a contribution from de-excitation nuclear lines. We found that the inclusion of a 2.223 MeV broad neutron capture line and the nuclear de-excitation lines did not improve the fit significantly and that the best-fit model for the emission observed by *Fermi*-GBM is a single PL.

As described in Section 5.2, the thin target is most likely, in which case following Equation (3) for the nonrelativistic case (<100 keV) the index of the electron number density $\delta \sim 2.7$. The radio spectrum observed by RSTN, shown in Figure 12, indicates microwave flux of 60 solar flux units (6×10^5 Jy), which peaks at frequency $\nu \sim 5$ GHz and declines as $\nu^{-\alpha}$ with the index $\alpha \sim 2$ above this frequency. Based on the relations described in Equation (4), we find that the required power-law index for the number density of electrons to be $\delta \sim 5$ or $\delta \sim 3.5$ in the relativistic or semirelativistic regime, respectively.

These values are steeper than the index $\delta \sim 2.7$ required in the nonrelativistic HXR range, indicating a steepening of the electron spectrum above 1 MeV. We therefore conclude that the GBM flux in the energy range 300 keV–10 MeV is probably not due to the accelerated electrons and no extra component is expected.

Sep14: The best-fit model for the combined HXR *Fermi*-GBM gamma-ray emission from the Sep14 flare is a single PL with an exponential decay with photon index 2.06 ± 0.01 and a cutoff energy of 90 ± 7 MeV. This is an extremely hard spectrum for a solar flare and if it is produced by a thin-target electron bremsstrahlung model as described above, it would require an electron number spectrum with index $\delta \sim 1.5$ in the nonrelativistic range steepening to an index $\delta \sim 2$ in the relativistic range. However, as noted above, inclusion of the 2.223 MeV line improves the fit slightly, indicating a possible contribution from nuclear de-excitation lines so the steepening of the electron spectrum could be larger.

The radio spectrum of this flare, shown in Figure 12, is also harder. The emission detected by the San Vito station has an index $\alpha \sim 1$ above 1 GHz. Again assuming optically thin synchrotron emission would require semirelativistic electrons with electron power-law index $\delta \sim 2.5$, indicating a slightly larger steepening at high energies. If the magnetic field is lower, which could be the case because of the larger loop, the steepening would be even larger ($\delta \sim 3$). Thus, for the Sep14

flare a power-law spectrum, steepening at higher energies, is required to explain the HXR and microwave observations from the loop-top region. This, plus a pion-decay component, provides an acceptable fit to the *Fermi* data from 30 keV to several GeV.

5.4. Acceleration Site and Mechanism

The above interpretation of the HXR and MWs indicates that the electrons responsible for these emissions are most likely accelerated in the reconnection region above the loop-top source either by turbulence (Petrosian 2012) or in merging islands manifested in the particle-in-cell simulations (Drake & Swisdak 2012). The electrons are trapped in the loop-top region long enough (longer than the crossing time of a fraction of second) to produce detectable bremsstrahlung and synchrotron radiation, but most of their energy is lost in thick-target footpoints hidden from near-Earth instruments because of high optical depths (Pesce-Rollins et al. 2015). This scenario can also explain the HXR emission from a BTL flare reported by Krucker et al. (2007). The steepening of the electron spectrum in the semirelativistic regime is due to a decrease in the acceleration rate and not to transport or energy-loss effects.

If the *Fermi*-LAT emission is from a thick-target photospheric site, then this emission site is located on the visible part of the disk. It cannot come from the occulted AR. The fact that accelerated protons reach the on-disk emission site provides strong evidence that the acceleration site is the CME environment, as suggested by Cliver et al. (1993) and Pesce-Rollins et al. (2015). They would diffuse across a wide range of magnetic fields, some connected to the AR behind the limb and some to the visible side of the Sun. These ions will most likely come from the downstream region of the CME shock while SEPs escape the upstream region. This can account for the differences in their spectral and temporal characteristics. It is then possible that the discrepancy between the spectral characteristics in 1–10 MeV for the GBM and radio observations can be explained by lower-energy 1–100 MeV protons or electrons precipitating to the visible disk side from the downstream region and producing the 1–10 MeV nuclear de-excitation line or bremsstrahlung emission below the chromosphere. These and other possibilities, for example, the production of *Fermi*-LAT emission by relativistic electrons either in the CME shock, in reconnection regions on current sheets behind the CME, or trapped in a large loop with strong convergence, will be discussed in future papers.

The *Fermi*-LAT Collaboration acknowledges support from a number of agencies and institutes for both development and the operation of the LAT as well as scientific data analysis. These

include NASA and DOE in the United States, CEA/Irfu and IN2P3/CNRS in France, ASI and INFN in Italy, MEXT, KEK, and JAXA in Japan, and the K. A. Wallenberg Foundation, the Swedish Research Council, and the National Space Board in Sweden. Additional support from INAF in Italy and CNES in France for science analysis during the operations phase is also gratefully acknowledged. W.L. was supported by NASA HGI grant NNX16AF78G and LWS grant NNX14AJ49G. V.P., W.L., and F.R.d.C. are supported by NASA grants NNX14AG03G, NNX13AF79G, and NNX12AO70G. V.P. and L.K. are grateful for RFBR grant 15-02-03717. We thank Nariaka Nitta and Meng Jin for fruitful discussions on SEPs and magnetic connectivity.

References

- Ackermann, M., Ajello, M., Albert, A., et al. 2014, *ApJ*, **787**, 15
 Ackermann, M., Ajello, M., Allafort, A., et al. 2012, *ApJ*, **745**, 144
 Ajello, M. A. A., Allafort, A., Baldini, L., et al. 2014, *ApJ*, **789**, 20
 Aptekar, R. L., Frederiks, D. D., Golenetskii, S. V., et al. 1995, *SSRv*, **71**, 265
 Arnaud, K. A. 1996, in ASP Conf. Ser. 101, *Astronomical Data Analysis Software and Systems V*, ed. G. H. Jacoby & J. Barnes (San Francisco, CA: ASP), 17
 Atwood, W. B., Abdo, A. A., Ackermann, M., et al. 2009, *ApJ*, **697**, 1071
 Barat, C., Trotter, G., Vilmer, N., et al. 1994, *ApJL*, **425**, L109
 Chen, Q., & Petrosian, V. 2013, *ApJ*, **777**, 33
 Cliver, E. W., Kahler, S. W., & Vestrand, W. T. 1993, *ICRC*, **3**, 91
 Debrunner, H., Flückiger, E., Grädel, H., Lockwood, J. A., & McGuire, R. E. 1988, *JGRA*, **93**, 7206
 Drake, J. F., & Swisdak, M. 2012, in APS April Meeting Abstracts, 3-D Particle-in-cell Simulations of the Sawtooth Crash, (Altana, GA), K1.022
 Dulk, G. A. 1985, *ARA&A*, **23**, 169
 Fitzpatrick, G., McBreen, S., Connaughton, V., & Briggs, M. 2012, *Proc. SPIE*, **8443**, 3
 Hurford, G. J., Schmahl, E. J., Schwartz, R. A., et al. 2002, *SoPh*, **210**, 61
 Krucker, S., White, S. M., & Lin, R. P. 2007, *ApJL*, **669**, L49
 Lin, R. P., & Hudson, H. S. 1971, *SoPh*, **17**, 412
 Malyshkin, L., & Kulsrud, R. 2001, *ApJ*, **549**, 402
 Mattox, J. R., Bertsch, D. L., Chiang, J., et al. 1996, *ApJ*, **461**, 396
 McTiernan, J. M., & Petrosian, V. 1990, *ApJ*, **359**, 541
 Meegan, C., Lichti, G., Bhat, P. N., et al. 2009, *ApJ*, **702**, 791
 Murphy, R. J., Dermer, C. D., & Ramaty, R. 1987, *ApJS*, **63**, 721
 Murphy, R. J., Kozlovsky, B., Kiener, J., & Share, G. H. 2009, *ApJS*, **183**, 142
 Nakajima, H., Nishio, M., Enome, S., et al. 1994, *IEEEP*, **82**, 705
 Nitta, N. V., Aschwanden, M. J., Boerner, P. F., et al. 2013, *SoPh*, **288**, 241
 Park, B. T., Petrosian, V., & Schwartz, R. A. 1997, *ApJ*, **489**, 358
 Pesce-Rollins, M., Omodei, N., Petrosian, V., et al. 2015, *ApJL*, **805**, L15
 Petrosian, V. 2012, *SSRv*, **173**, 535
 Petrosian, V. 2016, *ApJ*, **830**, 28
 Rybicki, G. B., & Lightman, A. P. 1979, *Radiative processes in astrophysics* (New York: Wiley-Interscience), 393
 Saldanha, R., Krucker, S., & Lin, R. P. 2008, *ApJ*, **673**, 1169
 Thakur, N., Gopalswamy, N., Xie, H., et al. 2014, *ApJL*, **790**, L13
 Thompson, W. T., & Wei, K. 2010, *SoPh*, **261**, 215
 Vestrand, W. T., & Forrest, D. J. 1993, *ApJL*, **409**, L69
 Vilmer, N., Trotter, G., Barat, C., et al. 1999, *A&A*, **342**, 575
 Wilks, S. S. 1938, *Ann. Math. Stat.*, **9**, 60