



**HAL**  
open science

# Chiral exact relations for helicities in Hall magnetohydrodynamic turbulence

Supratik Banerjee, Sébastien Galtier

► **To cite this version:**

Supratik Banerjee, Sébastien Galtier. Chiral exact relations for helicities in Hall magnetohydrodynamic turbulence. *Physical Review E*, 2016, 93, pp.033120. 10.1103/PhysRevE.93.033120. hal-01551972

**HAL Id: hal-01551972**

**<https://hal.science/hal-01551972>**

Submitted on 4 May 2022

**HAL** is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

# Chiral exact relations for helicities in Hall magnetohydrodynamic turbulence

Supratik Banerjee<sup>1,\*</sup> and Sébastien Galtier<sup>2,3,†</sup><sup>1</sup>*Institut für Geophysik und Meteorologie, Universität zu Köln, Pohlstrasse 3, D-50969 Köln, Germany*<sup>2</sup>*LPP, École polytechnique, F-91128 Palaiseau Cedex, France*<sup>3</sup>*Département de Physique, Université Paris-Sud, Orsay, France*

(Received 26 October 2015; revised manuscript received 1 February 2016; published 21 March 2016)

Besides total energy, three-dimensional incompressible Hall magnetohydrodynamics (MHD) possesses two inviscid invariants, which are the magnetic helicity and the generalized helicity. Exact relations are derived for homogeneous (nonisotropic) stationary Hall MHD turbulence (and also for its inertialess electron MHD limit) with nonzero helicities and in the asymptotic limit of large Reynolds numbers. The universal laws are written only in terms of mixed second-order structure functions, i.e., the scalar product of two different increments. It provides, therefore, a direct measurement of the dissipation rates for the corresponding invariant flux. This study shows that the generalized helicity cascade is strongly linked to the left polarized fluctuations, while the magnetic helicity cascade is linked to the right polarized fluctuations.

DOI: [10.1103/PhysRevE.93.033120](https://doi.org/10.1103/PhysRevE.93.033120)

## I. INTRODUCTION

Hall magnetohydrodynamics (MHD) is a monofluid plasma model appropriate for probing some of the physical processes (other than pure kinetic effects) at length scales smaller than the scales of standard MHD. Unlike MHD, this model introduces a decoupling between the electrons and the ions via the so-called Hall term in the generalized Ohm's law. The Hall effect becomes relevant when the typical length scales are smaller than the ion inertial length  $d_i$  ( $d_i \equiv c/\omega_{pi}$ , with  $c$  being the speed of light and  $\omega_{pi}$  being the ion plasma frequency). In a plasma for which the Alfvén mode controls the dynamics (incompressible or a collisionless plasma [1]), one can associate this length scale to the ion cyclotron angular frequency  $\omega_{ci} = V_A/d_i$  (with  $V_A$  being the Alfvén speed), and therefore, the Hall MHD regime is valid for the time scales of the order of or shorter than the ion cyclotron period  $2\pi/\omega_{ci}$ .

In space physics, the potential importance of Hall MHD is recognized for a range of phenomena varying from collisionless reconnection and disruption of Alfvénic filaments to small-scale solar wind turbulence [2–10]. Despite this fact, fewer analytical studies have been performed in Hall MHD [11–14], while the standard MHD has remained the subject of active theoretical and numerical research for more than three decades, especially in the framework of turbulence (see, e.g., Refs. [15–24]). A few direct numerical simulations, however, have been performed in Hall MHD turbulence to study either in a classical way the behavior of the velocity and magnetic field fluctuations [2,7,25–27] or in a more sophisticated way the separation of the fluctuations in terms of polarities (left and right) [28]. Shell models have also been used to investigate in particular the transition from standard to Hall MHD [29–31].

The Hall MHD model gets much simplified in the limit of small length scales ( $kd_i \gg 1$ ), where the ions are considered a motionless neutralizing background and the electron flow practically determines the electric current. This particular

limit of Hall MHD, called (the inertialess) electron MHD, is much easier to analyze, especially in the case of turbulence [32–35]. However, despite being a simple mathematical limit of Hall MHD, electron MHD is physically different from the former because the corresponding plasma dynamics is not governed by the force balance [36]. Besides having important roles in laser plasmas [37] and in magnetic field reconnection [38,39], electron MHD is also relevant for the understanding of the magnetic fluctuations in the collisionless solar wind at subion scales (see, e.g., Ref. [40]). Direct numerical simulations of isotropic electron MHD turbulence show that the magnetic energy spectrum scales like  $k^{-7/3}$  [41,42]. A plausible explanation for this spectrum has been provided by a heuristic model à la Kolmogorov [41], which turns out to be dimensionally compatible with an exact relation derived for third-order correlation functions [43]. However, the final form of this exact relation has not been reduced exclusively in terms of two-point fluctuations, which weakens any spectral prediction for electron MHD turbulence.

In addition to the total energy, Hall MHD permits two other inviscid invariants: the magnetic helicity  $H_M = \mathbf{a} \cdot \mathbf{b}$ , where  $\mathbf{b}$  is the magnetic field (which will be normalized to a velocity hereafter) and  $\mathbf{a}$  is its (normalized) vector potential, and the so-called generalized helicity  $H_G = (\mathbf{a} + d_i \mathbf{v}) \cdot (\mathbf{b} + d_i \mathbf{w})$ , where  $\mathbf{v}$  is the velocity and  $\mathbf{w} = \nabla \times \mathbf{v}$  is the vorticity vector [11,44–47]. In the electron MHD limit, we have  $\mathbf{v} = 0$ , which denotes  $H_M \equiv H_G$ . Hence, electron MHD allows two invariants: the magnetic energy (which is equivalent to its total energy) and the magnetic helicity. The role of the magnetic helicity in electron MHD turbulence was investigated recently with three-dimensional direct numerical simulations (with a mean magnetic field), which revealed that the propagation of one wave packet moving in one direction leads to the energy transfer towards larger scales [48,49]. This effect, interpreted as an inverse cascade, shows that one dispersive wave packet may produce another wave packet moving in the opposite direction as a result of the magnetic helicity conservation, which in turn leads to an inverse cascade. The impact of the magnetic helicity was investigated analytically in incompressible MHD [50]. An isotropic Kolmogorov-like exact relation was derived, but the final exact relation cannot

\*supratik.banerjee@uni-koeln.de

†sebastien.galtier@lpp.polytechnique.fr

be expressed only in terms of two-point fluctuations, thereby rendering the possibility of turbulent spectral prediction less evident. The role of the magnetic and generalized helicities in Hall MHD turbulence has never been studied analytically in great detail, although it is believed to be important, e.g., for the study of the plasma dynamo [5] for which a novel experiment has been designed recently [51].

In this paper, we derive an atypical class of exact relations corresponding to the magnetic helicity and generalized helicity conservation in homogeneous Hall MHD turbulence (isotropy is not assumed). In Sec. II we present the Hall MHD equations; in Sec. III we derive the exact law for the magnetic helicity cascade, while in Sec. IV the derivation is made for the generalized helicity cascade. In Sec. V we discuss the chiral nature of the helicity cascades, and we conclude the paper with a discussion in the last section.

## II. INCOMPRESSIBLE HALL MHD EQUATIONS

The magnetic helicity represents a quantitative measure of the self-linkage of the magnetic field; likewise, the generalized helicity represents a quantitative measure of the self-linkage of the generalized vorticity  $\Omega = \mathbf{b} + d_i \mathbf{w}$ . Below, we shall derive exact relations for average two-point fluctuations (i.e., increments) corresponding to the conservation of the magnetic helicity and generalized helicity in incompressible Hall MHD, for which the basic ideal and inviscid equations are

$$\partial_t \mathbf{v} = -\mathbf{v} \cdot \nabla \mathbf{v} - \nabla P + \mathbf{j} \times \mathbf{b} + \mathbf{f}_v, \quad (1)$$

$$\partial_t \mathbf{b} = \nabla \times (\mathbf{v} \times \mathbf{b}) - d_i \nabla \times (\mathbf{j} \times \mathbf{b}) + \mathbf{f}_b, \quad (2)$$

$$\nabla \cdot \mathbf{v} = 0, \quad (3)$$

$$\nabla \cdot \mathbf{b} = 0, \quad (4)$$

where  $P$  is the fluid pressure,  $\mathbf{j} = (\nabla \times \mathbf{b})$  denotes the current density, and  $\mathbf{f}_v$  and  $\mathbf{f}_b$  are stationary forcing terms (further properties will be given below). Equations (1) and (2) can easily be transformed into

$$\partial_t \mathbf{w} = \nabla \times [\mathbf{v} \times \mathbf{w} + \mathbf{j} \times \mathbf{b}] + \mathbf{f}_w, \quad (5)$$

$$\partial_t \mathbf{a} = (\mathbf{v} - d_i \mathbf{j}) \times \mathbf{b} + 2\nabla \psi + \mathbf{f}_a, \quad (6)$$

where  $\mathbf{f}_w$  and  $\mathbf{f}_a$  are the forcing terms and  $2\nabla \psi$  corresponds to an arbitrary choice of gauge.

In the previous equations the dissipative terms are not included. Generally, in MHD we introduce a Laplacian operator for both the velocity and magnetic field equations in order to mimic in a simple way the mechanism of dissipation which involves kinetic effects [52]. In MHD turbulence these terms are fundamental to ensure the existence of a well-resolved inertial range in which the dissipation (and forcing) is negligible. In Hall MHD turbulence this modeling is less relevant, especially for the magnetic field, because the nonlinear Hall term involves a derivative of order 2 and because the magnetic spectrum is generally steeper than in MHD. Then, with a classical Laplacian the dissipation is generally less localized at small scales than in MHD, which may reduce the size of the inertial range. This means that to get an inertial

range well separated from the small-scale dissipative range it is better to have a dissipative term with an order of derivative higher than 2 (e.g., a bi-Laplacian).

A detailed discussion of the dissipation mechanism is beyond the scope of this paper because (i) it depends on the application from which our derivation is independent and (ii) there are a variety of possible sources. However, it is relevant to make a comparison with the solar wind, on which a lot of studies are currently focused. The weakly collisional plasma conditions in the solar wind imply that the mechanism responsible for the dissipation of turbulent fluctuations is likely to be of a kinetic nature like the damping of linear waves [53]. This damping corresponds certainly to a term different from a Laplacian or double Laplacian. If the damping is either localized at some scales or significantly smaller in amplitude than the nonlinear fluctuations, then it will not affect significantly the inertial range and therefore our conclusion. Note, however, that the question of the relative importance of the damping is currently being actively debated. For example, in a more recent paper [54] it is claimed that the solar wind data are compatible with the conclusion that the nonlinear time scale is comparable to or shorter than the typical time scale of wave damping, which renders the linear treatment questionable.

The other important term for this type of analysis is the external forcing. In our case, we are mainly interested in the magnetic and generalized helicities. The injection of magnetic helicity is evoked in different astrophysical contexts like the sun, where the source can be the shearing motions [55]. It is also discussed in the context of the dynamo problem, where an inverse cascade of helicity can produce (or regenerate) a large-scale magnetic field [15,56]. Magnetic turbulence in astrophysical systems is often governed by two effects, buoyancy and rotation, which naturally lead to helical flows and twisted field lines. The result is the production of a net magnetic helicity and also probably some generalized helicity (since it contains the magnetic and kinetic helicities). This forcing is often made in a limited range of scales. For example, in the case of the geodynamo the typical scale ( $\sim 10^3$  km) corresponds roughly to the size (a fraction of the thickness) of the convection layer.

## III. EXACT RELATION FOR MAGNETIC HELICITY CONSERVATION

When a flux of magnetic helicity is injected into a turbulent plasma, at a typical wave number  $k_f$ , an inverse cascade is expected. This happens only at the MHD scale if  $k_f d_i < 1$  [15] or at subionic Hall MHD scales if  $k_f d_i > 1$  [49,57], with eventually the possibility to extend the inverse cascade to MHD scales. In this section, we shall consider a forcing at intermediate scale such that  $k_f d_i > 1$  and focus the analysis on the inertial range at a scale  $k \ll k_f$  (small-scale forcing). We see immediately a potential problem to derive a (statistically) stationary law in a finite size system: we need a dissipation at large-scale to counterbalance the injection of magnetic helicity and avoid the formation of a condensate. In two-dimensional hydrodynamic turbulence a similar situation is found but with an inverse cascade of kinetic energy [58]. Frictional dissipation, usually due to, e.g., friction between

the fluid and substrate, is then introduced into the system either numerically and/or analytically to avoid the formation of a condensate [59]. We will follow this line and assume the presence of some large-scale dissipation whose origin may depend on the problem (e.g., the wall in a dynamo experiment).

Note that in the case of the solar wind, since Hall MHD contains the MHD scales, an inverse cascade of magnetic helicity can continue potentially up to the largest scales of the system, i.e., to frequencies much lower than the cyclotron frequency. However, another limitation exists: the magnetic helicity is not conserved in MHD if a uniform magnetic field is present (see, e.g., Ref. [60]), and the solar wind is composed of a large-scale magnetic field plus fluctuations. In other words, when the inverse cascade crosses the scale  $\sim d_i$ , the transfer gets worse, and we have therefore a nonlinear (nondissipative) limit in length scale to the inverse cascade.

Unlike [50], we define the *symmetric* two-point correlation function associated with the magnetic helicity as

$$R_H = R'_H = \left\langle \frac{\mathbf{a} \cdot \mathbf{b}' + \mathbf{a}' \cdot \mathbf{b}}{2} \right\rangle, \quad (7)$$

where the primed and unprimed quantities correspond to points  $\mathbf{x}'$  and  $\mathbf{x}$ , respectively, with  $\mathbf{x}' = \mathbf{x} + \mathbf{r}$ , and  $\langle \cdot \rangle$  means an ensemble average (which is equivalent by ergodicity to a spatial average in homogeneous turbulence). Using Eqs. (2) and (6) and defining  $\mathbf{U} = \mathbf{v} - d_i \mathbf{j}$ , we can write

$$\begin{aligned} \partial_t (R_H + R'_H) &= \langle \mathbf{a}' \cdot \partial_t \mathbf{b} + \mathbf{b} \cdot \partial_t \mathbf{a}' + \mathbf{a} \cdot \partial_t \mathbf{b}' + \mathbf{b}' \cdot \partial_t \mathbf{a} \rangle \\ &= \langle \mathbf{a}' \cdot [\nabla \times (\mathbf{U} \times \mathbf{b})] + \mathbf{b} \cdot (\mathbf{U}' \times \mathbf{b}') + 2\mathbf{b} \cdot \nabla' \psi' \rangle + \mathcal{D}_H \\ &\quad + \langle \mathbf{a} \cdot [\nabla' \times (\mathbf{U}' \times \mathbf{b}')] + \mathbf{b}' \cdot (\mathbf{U} \times \mathbf{b}) + 2\mathbf{b}' \cdot \nabla \psi \rangle + \mathcal{F}_H \\ &= \langle \nabla \cdot [(\mathbf{U} \times \mathbf{b}) \times \mathbf{a}'] + \mathbf{b} \cdot [\mathbf{U}' \times \mathbf{b}' + 2\nabla' \psi'] \rangle + \mathcal{D}_H \\ &\quad + \langle \nabla' \cdot [(\mathbf{U}' \times \mathbf{b}') \times \mathbf{a}] + \mathbf{b}' \cdot [\mathbf{U} \times \mathbf{b} + 2\nabla \psi] \rangle + \mathcal{F}_H, \end{aligned} \quad (8)$$

with  $\mathcal{D}_H$  being large-scale dissipation and

$$\mathcal{F}_H = \langle \mathbf{a}' \cdot \mathbf{f}_b + \mathbf{b} \cdot \mathbf{f}'_a + \mathbf{a} \cdot \mathbf{f}'_b + \mathbf{b}' \cdot \mathbf{f}_a \rangle \quad (9)$$

being small-scale forcing. By using the relation  $\mathbf{b} = \nabla \times \mathbf{a}$  and the statistical homogeneity, we obtain

$$\begin{aligned} &\langle \nabla \cdot [(\mathbf{U} \times \mathbf{b}) \times \mathbf{a}'] + \nabla' \cdot [(\mathbf{U}' \times \mathbf{b}') \times \mathbf{a}] \rangle \\ &= -\langle \nabla' \cdot [(\mathbf{U} \times \mathbf{b}) \times \mathbf{a}'] + \nabla \cdot [(\mathbf{U}' \times \mathbf{b}') \times \mathbf{a}] \rangle \\ &= \langle (\mathbf{U} \times \mathbf{b}) \cdot \mathbf{b}' + (\mathbf{U}' \times \mathbf{b}') \cdot \mathbf{b} \rangle. \end{aligned} \quad (10)$$

Inserting this above simplification in Eq. (8), we find

$$\begin{aligned} \partial_t \left( \frac{R_H + R'_H}{2} \right) &= \langle [(\mathbf{U} \times \mathbf{b}) + \nabla \psi] \cdot \mathbf{b}' + [(\mathbf{U}' \times \mathbf{b}') + \nabla' \psi'] \cdot \mathbf{b} \rangle \\ &\quad + \frac{\mathcal{D}_H}{2} + \frac{\mathcal{F}_H}{2} \\ &= -\langle \delta(\mathbf{U} \times \mathbf{b}) \cdot \delta \mathbf{b} \rangle + \frac{\mathcal{D}_H}{2} + \frac{\mathcal{F}_H}{2}, \end{aligned} \quad (11)$$

where we use

$$\langle \nabla \psi \cdot \mathbf{b}' + \nabla' \psi' \cdot \mathbf{b} \rangle = -\langle \psi(\nabla' \cdot \mathbf{b}') + \psi'(\nabla \cdot \mathbf{b}) \rangle = 0$$

due to statistical homogeneity. In the final step, we consider a stationary state (in the limit of infinite kinetic and magnetic Reynolds numbers) corresponding to the magnetic helicity conservation for which the left-hand side of Eq. (11) vanishes. Under this condition, we derive an expression for the inertial range within which the small-scale forcing term becomes negligible and the large-scale dissipative effect gives the mean magnetic helicity flux dissipation rate  $\eta_M$ , which is also equal to the mean magnetic helicity flux transfer rate. The final form of the exact relation is then given by

$$\eta_M = \langle \delta(\mathbf{U} \times \mathbf{b}) \cdot \delta \mathbf{b} \rangle, \quad (12)$$

which can also be written as

$$\eta_M = \langle \delta(\mathbf{v} \times \mathbf{b}) \cdot \delta \mathbf{b} \rangle - d_i \langle \delta(\mathbf{j} \times \mathbf{b}) \cdot \delta \mathbf{b} \rangle. \quad (13)$$

In the large-scale MHD limit ( $kd_i \ll 1$ ) we may recover the result obtained by [50] when it is symmetrized. In the small-scale electron MHD limit ( $kd_i \gg 1$ , or, equivalently,  $\mathbf{v} \rightarrow 0$ ) the above relationship is further simplified to

$$\eta_M = d_i \langle \delta(\mathbf{b} \times \mathbf{j}) \cdot \delta \mathbf{b} \rangle. \quad (14)$$

Equation (13) is the first main result of this paper. It is an exact relation valid for three-dimensional homogeneous Hall MHD turbulence without the assumption of isotropy. It is relevant to make a comparison with previous results and in particular with the magnetic helicity law derived in MHD [50], which is the large-scale limit of Hall MHD. Unlike [50], the assumption of isotropy is not made in our derivation, giving the law a broader application, such as to space plasmas, for which turbulence is rarely isotropic (see, e.g., Ref. [61]). Additionally, law (13) implies products of two increments (i.e., only fluctuations), a situation well adapted to turbulence, whereas the final law in [50] did not consist of two-point increments. A comparison with the classical derivation of the four-thirds laws for the energy (see, e.g., Ref. [62]) reveals another difference: our derivation does not lead to the appearance of a divergence operator, which renders difficult the evaluation of the energy cascade rate in general (i.e., when isotropy is not assumed). In our case, we see that the evaluation of  $\eta_M$  requires only the computation of a scalar product of two increments. This property should be helpful for the study of the magnetic helicity effects in space (anisotropic) plasmas, e.g., in the solar wind [63]. Finally, from expressions (13) and (14), one can easily infer that in both Hall MHD and electron MHD no turbulent flux of magnetic helicity exists if the system satisfies the corresponding Beltrami conditions, i.e.,  $\mathbf{U} \parallel \mathbf{b}$  and  $\mathbf{j} \parallel \mathbf{b}$ , respectively. Therefore, we recover well the theory of plasma relaxation to Beltrami alignment [11].

A power law spectrum  $\sim k^{-2}$  is expected for the magnetic helicity when an inverse cascade happens at MHD scales [15] (for simplicity, isotropy is assumed for the discussion). Dimensionally (assuming simply the maximal helicity state), this scaling corresponds to a  $-1$  power law for the energy and thus to a constant  $\delta b$ . The exact relation (13) tells us at the dimensional level (for simplicity we do not consider the effect of the scalar product, which may give an additional scaling dependence) that, necessarily, the other increment  $\delta(\mathbf{v} \times \mathbf{b})$  has no scale dependence to keep globally constant the right-hand side term. The situation is different at subion electron MHD

scales, where the inverse cascade of magnetic helicity may lead to [by using as before a Kolmogorov phenomenology with the assumption of a maximal helicity state but in which the transfer time is now  $\sim r^2/(d_i b)$ ] a corresponding power spectrum  $\sim k^{-8/3}$ . Hence, a  $-5/3$  magnetic energy spectrum is found, and therefore,  $\delta b \sim r^{1/3}$ . The exact relation (14) tells us that in this case, necessarily, we have dimensionally  $\delta(\mathbf{b} \times \mathbf{j}) \sim r^{-1/3}$ . A possible interpretation of this scaling law is that a phenomenon of alignment between  $\mathbf{b}$  and  $\mathbf{j}$  appears when one goes from small to large electron MHD scales. However, note that in two dimensions this conclusion may change [64].

#### IV. EXACT RELATION FOR GENERALIZED HELICITY CONSERVATION

In this section, we shall derive a universal law for the generalized helicity  $H_G$ . Hereinafter, we shall denote  $\mathbf{Y} = (\mathbf{a} + d_i \mathbf{v})$  as the generalized vector potential, whereas  $\mathbf{\Omega}$  is the generalized vorticity. The generalized vorticity  $\mathbf{\Omega}$  is somewhat analogous to the magnetic field in standard MHD [44]. Indeed, both quantities obey the same Lagrangian equation, which is, for  $\mathbf{\Omega}$ ,

$$\frac{d\mathbf{\Omega}}{dt} = \mathbf{\Omega} \cdot \nabla \mathbf{v}. \quad (15)$$

By applying Helmholtz's law to Hall MHD, we see that the generalized vorticity lines are frozen into the plasma. However, the fact that at small scales  $\mathbf{v} \times \mathbf{b} - d_i \mathbf{j} \times \mathbf{b} \simeq \mathbf{v}_e \times \mathbf{b}$  leads to

$$\frac{d\mathbf{b}}{dt} = \mathbf{b} \cdot \nabla \mathbf{v}_e, \quad (16)$$

which shows that the magnetic field, which is no longer frozen in the fluid due to the Hall term, is still frozen with respect to the electronic fluid. Like the magnetic helicity in incompressible MHD, it is possible to show that the generalized helicity of a generalized vorticity tube is conserved over time [45]. Therefore,  $H_G$  provides a measure of the degree of structural complexity (the topology) of an incompressible Hall MHD flow.

With our notation, the generalized helicity can be written as  $H_G = \mathbf{Y} \cdot \mathbf{\Omega}$ . Using Eqs. (1)–(6), we obtain

$$\partial_t \mathbf{Y} = \mathbf{v} \times \mathbf{\Omega} - \nabla P_G + \mathbf{f}_Y, \quad (17)$$

$$\partial_t \mathbf{\Omega} = \nabla \times (\mathbf{v} \times \mathbf{\Omega}) + \mathbf{f}_\Omega, \quad (18)$$

where  $P_G = d_i v^2/2 + d_i P - 2\psi$  is a generalized pressure and  $\mathbf{f}_{Y,\Omega}$  denote the corresponding forcing terms. (The dissipative terms will be introduced later.) Now we construct the symmetric two-point correlation function for the generalized helicity, namely,

$$R_G = R'_G = \left\langle \frac{\mathbf{\Omega} \cdot \mathbf{Y}' + \mathbf{\Omega}' \cdot \mathbf{Y}}{2} \right\rangle. \quad (19)$$

Below, we shall derive the evolution equation of the correlation function. For the same reasons as before we do not introduce dissipative terms. Only external forcing terms  $\mathcal{F}_G$  are introduced in the equations; then we get

$$\begin{aligned} \partial_t (R_G + R'_G) \\ = \langle [-\nabla P_c + \mathbf{v} \times \mathbf{\Omega}] \cdot \mathbf{\Omega}' + \mathbf{Y} \cdot [\nabla' \times (\mathbf{v}' \times \mathbf{\Omega}')] \rangle \end{aligned}$$

$$\begin{aligned} + \langle [-\nabla' P'_c + \mathbf{v}' \times \mathbf{\Omega}'] \cdot \mathbf{\Omega} + \mathbf{Y}' \cdot [\nabla \times (\mathbf{v} \times \mathbf{\Omega})] \rangle + \mathcal{F}_G \\ = \langle (\mathbf{v} \times \mathbf{\Omega}) \cdot \mathbf{\Omega}' + (\mathbf{v}' \times \mathbf{\Omega}') \cdot \mathbf{\Omega} \rangle + \langle \nabla' \cdot [(\mathbf{v}' \times \mathbf{\Omega}') \times \mathbf{Y}] \\ + \nabla \cdot [(\mathbf{v} \times \mathbf{\Omega}) \times \mathbf{Y}'] \rangle + \mathcal{F}_G, \end{aligned} \quad (20)$$

where  $\mathcal{F}_G = \langle \mathbf{\Omega} \cdot \mathbf{f}'_Y + \mathbf{Y}' \cdot \mathbf{f}_\Omega + \mathbf{\Omega}' \cdot \mathbf{f}_Y + \mathbf{Y} \cdot \mathbf{f}'_\Omega \rangle$ . The definition  $\mathbf{\Omega} = \nabla \times \mathbf{Y}$  and the assumption of statistical homogeneity lead to the relation

$$\begin{aligned} \langle \nabla' \cdot [(\mathbf{v}' \times \mathbf{\Omega}') \times \mathbf{Y}] + \nabla \cdot [(\mathbf{v} \times \mathbf{\Omega}) \times \mathbf{Y}'] \rangle \\ = -\langle \nabla \cdot [(\mathbf{v}' \times \mathbf{\Omega}') \times \mathbf{Y}] + \nabla' \cdot [(\mathbf{v} \times \mathbf{\Omega}) \times \mathbf{Y}'] \rangle \\ = \langle (\mathbf{v}' \times \mathbf{\Omega}') \cdot \mathbf{\Omega} + (\mathbf{v} \times \mathbf{\Omega}) \cdot \mathbf{\Omega}' \rangle. \end{aligned} \quad (21)$$

Substituting this expression into Eq. (20), we get

$$\begin{aligned} \partial_t \left( \frac{R_G + R'_G}{2} \right) &= \langle (\mathbf{v}' \times \mathbf{\Omega}') \cdot \mathbf{\Omega} + (\mathbf{v} \times \mathbf{\Omega}) \cdot \mathbf{\Omega}' \rangle + \frac{\mathcal{F}_G}{2} \\ &= -\langle \delta(\mathbf{v} \times \mathbf{\Omega}) \cdot \delta \mathbf{\Omega} \rangle + \frac{\mathcal{F}_G}{2}, \end{aligned} \quad (22)$$

as  $(\mathbf{v} \times \mathbf{\Omega}) \cdot \mathbf{\Omega} = (\mathbf{v}' \times \mathbf{\Omega}') \cdot \mathbf{\Omega}' = 0$ .

For the final step of the development we shall introduce a dissipative term which is necessary to make the assumption of a stationary state corresponding to the conservation of the generalized helicity. By construction, the generalized vorticity and the vector potential are made of two types of terms: a magnetic term and a kinetic term. That means we can distinguish two situations: the case where the magnetic energy  $\mathcal{E}^m$  is much greater than the kinetic energy  $\mathcal{E}^u$  and the opposite situation, i.e.,  $\mathcal{E}^u \gg \mathcal{E}^m$ . In the first case, the generalized helicity reduces to the magnetic helicity, and as explained in the previous section, we may expect an inverse cascade. This situation corresponds actually (at subion scales) to electron MHD. In the second case, the generalized helicity reduces to the kinetic helicity. In this case the Hall MHD equations can be simplified to the Navier-Stokes equations for which the kinetic helicity is indeed an invariant that leads to a direct cascade [65]. Therefore, we see that both a direct and an inverse cascade may happen for the generalized helicity. Note that this conclusion may also be reached when one uses a Gibbs ensemble analysis [66]. To establish an exact relation that is valid well inside the inertial zone we shall introduce a large- or small-scale dissipation which allows us to assume a stationary state. In practice, that corresponds to a small- or large-scale forcing, respectively. We introduce the mean generalized helicity flux dissipative rate  $\eta_G$  and obtain after simplification

$$\pm \eta_G = \langle \delta(\mathbf{v} \times \mathbf{\Omega}) \cdot \delta \mathbf{\Omega} \rangle, \quad (23)$$

where the plus and minus signs correspond to direct and inverse cascades, respectively (by definition  $\eta_G$  is positive).

Expression (23) is the second main result of this paper. It is an exact relation for three-dimensional homogeneous Hall MHD turbulence without the assumption of isotropy. When the magnetic energy dominates, we recover expression (13) both in the large-scale and small-scale limits (in the latter case  $\mathbf{v} - d_i \mathbf{j} \rightarrow \mathbf{v}_e$ ).

## V. CHIRALITY OF THE HELICITY CASCADES

Interestingly, each helicity cascade (magnetic and generalized) can be associated with a specific polarization. A simple way to recognize this fact is to write laws (13) and (23) for the helicities in terms of the following generalized vortices and velocities [11,28]:

$$\boldsymbol{\Omega}_L = \mathbf{b} + d_i \nabla \times \mathbf{v}, \quad (24)$$

$$\boldsymbol{\Omega}_R = \mathbf{b}, \quad (25)$$

$$\mathbf{v}_L = \mathbf{v}, \quad (26)$$

$$\mathbf{v}_R = \mathbf{v} - d_i \nabla \times \mathbf{b}, \quad (27)$$

where indices  $L$  and  $R$  refer to different polarities. To understand why these variables correspond to the  $L$  and  $R$  polarities, one possibility is to study the exclusive linear effect of each variable in the absence of the other [28]. The Hall MHD equations can be written as

$$\partial_t \boldsymbol{\Omega}_{L/R} = \nabla \times (\mathbf{v}_{L/R} \times \boldsymbol{\Omega}_{L/R}). \quad (28)$$

When the equations are linearized, by imposing a uniform magnetic field  $\mathbf{b}_0$ , we get

$$\partial_t \tilde{\boldsymbol{\Omega}}_{L/R} = \nabla \times (\tilde{\mathbf{v}}_{L/R} \times \mathbf{b}_0), \quad (29)$$

where  $\tilde{\mathbf{a}}$  denotes the linear (first-order) perturbation of variable  $\mathbf{a}$ . For the case where only the  $L$  variables survive, we have  $\tilde{\boldsymbol{\Omega}}_R = 0 = \tilde{\mathbf{v}}_R$ . Under this condition expression (29) reduces to

$$d_i \partial_t (\nabla \times \tilde{\mathbf{v}}) = \nabla \times (\tilde{\mathbf{v}} \times \mathbf{b}_0). \quad (30)$$

For every perturbation, we assume a plane wave solution of the form  $\exp[i(\mathbf{k} \cdot \mathbf{x} - \omega t)]$  with angular frequency  $\omega$  and wave vector  $\mathbf{k}$ . The above relation therefore transforms (with  $\partial_t \mapsto -i\omega$  and  $\nabla \mapsto i\mathbf{k}$ ) into

$$\omega d_i (\mathbf{k} \times \hat{\mathbf{v}}) = i b_0 k_{\parallel} \hat{\mathbf{v}}, \quad (31)$$

where  $\hat{\boldsymbol{\xi}}$  denotes the Fourier transform of variable  $\boldsymbol{\xi}$ . This equation represents a left-circularly polarized wave and also necessarily gives  $\omega = b_0 k_{\parallel} / (k d_i)$ , which identifies the mode as a known ion-cyclotronic mode. For the other case, we have similarly  $\tilde{\boldsymbol{\Omega}}_L = 0 = \tilde{\mathbf{v}}_L$ , and the linearization of Eq. (28) gives

$$\partial_t \tilde{\mathbf{b}} = -d_i \nabla \times [(\nabla \times \tilde{\mathbf{b}}) \times \mathbf{b}_0]. \quad (32)$$

After Fourier transformation, this simply becomes

$$i\omega \hat{\mathbf{b}} = -d_i b_0 k_{\parallel} (\mathbf{k} \times \hat{\mathbf{b}}), \quad (33)$$

which is the equation of a right-circularly polarized wave mode. In addition, we also get  $\omega = b_0 k_{\parallel} / k d_i$ , thereby recognizing the well-known right-circularly polarized whistler mode.

It is straightforward to see (by a simple substitution) that the exact relations (13) and (23) can be written only in terms of the generalized fields. We find

$$\eta_M = \langle \delta(\mathbf{v}_R \times \boldsymbol{\Omega}_R) \cdot \delta \boldsymbol{\Omega}_R \rangle, \quad (34)$$

$$\pm \eta_G = \langle \delta(\mathbf{v}_L \times \boldsymbol{\Omega}_L) \cdot \delta \boldsymbol{\Omega}_L \rangle. \quad (35)$$

The form of expressions (34) and (35) demonstrates that, on one hand, the magnetic helicity cascade is by nature a process implying only the right-handed fluctuations, whereas the generalized helicity cascade implies only the left-handed fluctuations. While the former property may be expected since the magnetic helicity is an ideal invariant of electron MHD, the latter is less trivial. Following a recent analysis [28], we may call ion MHD the regime where only right fluctuations persist.

## VI. DISCUSSION

Equations (13), (23), (34), and (35) are the main results of this paper. These exact relations are valid for homogeneous Hall MHD turbulence without the assumption of isotropy. Unlike the traditional Yaglom form, the nonisotropic expression does not involve any global divergence term and provides therefore a direct evaluation (using numerical simulations or observational data) of the transfer rates for the helicities. However, care should be taken because a stationary state corresponding to the helicity conservation is essential in order to verify these relations. In addition, we get helicity laws for Hall MHD which are purely expressible in terms of two-point increments. We have also shown that the laws for the helicities can be written in terms of the generalized vortices and velocities, which reveals the chiral properties of the helicity cascades, with the magnetic and generalized helicity cascades being associated with right- and left-handed fluctuations, respectively. In addition, expressions (34) and (35) show trivially that a Beltrami flow ( $\boldsymbol{\Omega}_i = a_i \mathbf{v}_i$ ), which is the state of maximum helicity, cannot produce a nonlinear cascade for the helicities. A summary of our results is proposed in Fig. 1: in this fluid scenario, we do not introduce kinetic effects that should appear especially around the scale  $k d_i \sim 1$ . Note, finally, that expressions (13) and (23) or (34) and (35) can be used

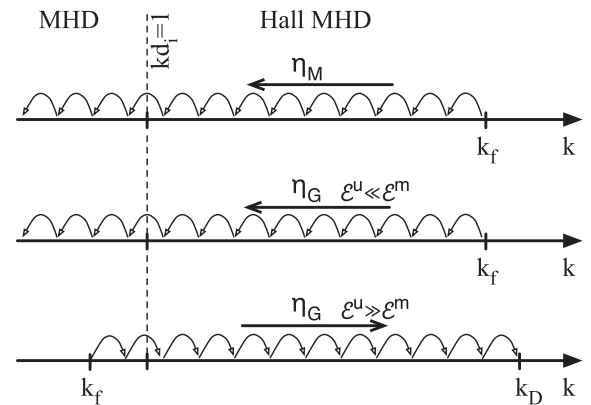


FIG. 1. Schematic view of the turbulent helicity cascades in Hall MHD. Top: inverse cascade of magnetic helicity with a small-scale forcing ( $k_f d_i \gg 1$ ). Middle: inverse cascade of generalized helicity in a magnetic regime ( $\mathcal{E}^m \gg \mathcal{E}^u$ ) and with a small-scale forcing ( $k_f d_i \gg 1$ ). Bottom: direct cascade of generalized helicity in a kinetic regime ( $\mathcal{E}^u \gg \mathcal{E}^m$ ) and with a large-scale forcing ( $k_f d_i < 1$ ). The vertical dashed line ( $k d_i = 1$ ) separates the MHD scales (on the left) from the Hall MHD scales (on the right).  $k_D^{-1}$  represents the dissipation length scale.

to evaluate experimentally the associated cascade rates. For example, the recently launched Magnetospheric Multiscale Mission (MMS) composed of four identically instrumented

spacecraft may provide this information during incursions into the solar wind by measuring the different increments at subion scales (for both magnetic and plasma data).

- 
- [1] J. W. Belcher and L. Davis, Jr., *J. Geophys. Res.* **76**, 3534 (1971).
- [2] S. Ghosh and M. L. Goldstein, *J. Plasma Phys.* **57**, 129 (1997).
- [3] A. Bhattacharjee, *Annu. Rev. Astron. Astrophys.* **42**, 365 (2004).
- [4] J. Dreher, D. Laveder, R. Grauer, T. Passot, and P. L. Sulem, *Phys. Plasmas* **12**, 052319 (2005).
- [5] P. D. Mininni, D. O. Gómez, and S. M. Mahajan, *Astrophys. J.* **619**, 1019 (2005).
- [6] S. Galtier, *J. Low Temp. Phys.* **145**, 59 (2006).
- [7] P. D. Mininni, A. Alexakis, and A. Pouquet, *J. Plasma Phys.* **73**, 377 (2007).
- [8] F. Sahrhoui, S. Galtier, and G. Belmont, *J. Plasma Phys.* **73**, 723 (2007).
- [9] O. Alexandrova, J. Saur, C. Lacombe, A. Mangeney, J. Mitchell, S. J. Schwartz, and P. Robert, *Phys. Rev. Lett.* **103**, 165003 (2009).
- [10] K. H. Kiyani, S. C. Chapman, Y. V. Khotyaintsev, M. W. Dunlop, and F. Sahrhoui, *Phys. Rev. Lett.* **103**, 075006 (2009).
- [11] Z. Yoshida and S. M. Mahajan, *Phys. Rev. Lett.* **88**, 095001 (2002).
- [12] S. Galtier, *Phys. Rev. E* **77**, 015302 (2008).
- [13] K. Araki, *J. Phys. A* **48**, 175501 (2015).
- [14] K. Araki, *Phys. Rev. E* **92**, 063106 (2015).
- [15] A. Pouquet, U. Frisch, and J. Leorat, *J. Fluid Mech.* **77**, 321 (1976).
- [16] S. Galtier, S. V. Nazarenko, A. C. Newell, and A. Pouquet, *J. Plasma Phys.* **63**, 447 (2000).
- [17] W.-C. Müller, D. Biskamp, and R. Grappin, *Phys. Rev. E* **67**, 066302 (2003).
- [18] E. Buchlin, V. Aletti, S. Galtier, M. Velli, G. Einaudi, and J.-C. Vial, *Astron. Astrophys.* **406**, 1061 (2003).
- [19] A. Alexakis and Y. Ponty, *Phys. Rev. E* **77**, 056308 (2008).
- [20] W. H. Matthaeus, S. Oughton, and Y. Zhou, *Phys. Rev. E* **79**, 035401 (2009).
- [21] S. Galtier, *Astrophys. J.* **746**, 184 (2012).
- [22] A. Beresnyak, *Phys. Rev. Lett.* **106**, 075001 (2011).
- [23] S. Banerjee and S. Galtier, *Phys. Rev. E* **87**, 013019 (2013).
- [24] R. Meyrand, K. H. Kiyani, and S. Galtier, *J. Fluid Mech.* **770**, R1 (2015).
- [25] D. Shaikh and P. K. Shukla, *Phys. Rev. Lett.* **102**, 045004 (2009).
- [26] H. Miura and K. Araki, *Phys. Plasmas* **21**, 072313 (2014).
- [27] J. E. Stawarz and A. Pouquet, *Phys. Rev. E* **92**, 063102 (2015).
- [28] R. Meyrand and S. Galtier, *Phys. Rev. Lett.* **109**, 194501 (2012).
- [29] S. Galtier and E. Buchlin, *Astrophys. J.* **656**, 560 (2007).
- [30] D. Hori and H. Miura, *Plasma Fusion Res.* **3**, S1053 (2008).
- [31] D. Banerjee, S. S. Ray, G. Sahoo, and R. Pandit, *Phys. Rev. Lett.* **111**, 174501 (2013).
- [32] A. Kingsep, K. Chukbar, and V. Yankov, *Rev. Plasma Phys.* **16**, 243 (1990).
- [33] P. Goldreich and A. Reisenegger, *Astrophys. J.* **395**, 250 (1992).
- [34] R. Meyrand and S. Galtier, *Astrophys. J.* **721**, 1421 (2010).
- [35] P. H. Diamond, A. Hasegawa, and K. Mima, *Plasma Phys. Controlled Fusion* **53**, 124001 (2011).
- [36] M. Lyutikov, *Phys. Rev. E* **88**, 053103 (2013).
- [37] Y. Sentoku, K. Mima, P. Kaw, and K. Nishikawa, *Phys. Rev. Lett.* **90**, 155001 (2003).
- [38] S. V. Bulanov, F. Pegoraro, and A. S. Sakharov, *Phys. Fluids B* **4**, 2499 (1992).
- [39] J. F. Drake, R. G. Kleva, and M. E. Mandt, *Phys. Rev. Lett.* **73**, 1251 (1994).
- [40] R. Meyrand and S. Galtier, *Phys. Rev. Lett.* **111**, 264501 (2013).
- [41] D. Biskamp, E. Schwarz, and J. F. Drake, *Phys. Rev. Lett.* **76**, 1264 (1996).
- [42] C. S. Ng, A. Bhattacharjee, K. Germaschewski, and S. Galtier, *Phys. Plasmas* **10**, 1954 (2003).
- [43] S. Galtier, *J. Geophys. Res.* **113**, 1102 (2008).
- [44] L. Woltjer, *Proc. Natl. Acad. Sci. USA* **44**, 489 (1958).
- [45] L. Turner, *IEEE Trans. Plasma Sci.* **14**, 849 (1986).
- [46] V. Krishan and S. M. Mahajan, *J. Geophys. Res.* **109**, A11105 (2004).
- [47] S. Galtier, *J. Plasma Phys.* **72**, 721 (2006).
- [48] J. Cho, *Phys. Rev. Lett.* **106**, 191104 (2011).
- [49] H. Kim and J. Cho, *Astrophys. J.* **801**, 75 (2015).
- [50] H. Politano, T. Gomez, and A. Pouquet, *Phys. Rev. E* **68**, 026315 (2003).
- [51] C. M. Cooper, J. Wallace, M. Brookhart, M. Clark, C. Collins, W. X. Ding, K. Flanagan, I. Khalzov, Y. Li, J. Milhone *et al.*, *Phys. Plasmas* **21**, 013505 (2014).
- [52] G. Belmont, R. Grappin, F. Mottez, F. Pantellini, and G. Pelletier, *Collisionless Plasmas in Astrophysics* (Wiley, Hoboken, NJ, 2013).
- [53] J. M. TenBarge, G. G. Howes, and W. Dorland, *Astrophys. J.* **774**, 139 (2013).
- [54] W. H. Matthaeus, S. Oughton, K. T. Osman, S. Servidio, M. Wan, S. P. Gary, M. A. Shay, F. Valentini, V. Roytershteyn, H. Karimabadi *et al.*, *Astrophys. J.* **790**, 155 (2014).
- [55] P. Démoulin, C. H. Mandrini, L. Van Driel-Gesztelyi, M. C. Lopez Fuentes, and G. Aulanier, *Sol. Phys.* **207**, 87 (2002).
- [56] D. Biskamp, *Magnetohydrodynamic Turbulence* (Cambridge University Press, Cambridge, 2003).
- [57] S. Galtier and R. Meyrand, *J. Plasma Phys.* **81**, 3206 (2015).
- [58] R. H. Kraichnan, *Phys. Fluids* **10**, 1417 (1967).
- [59] M. Chertkov, C. Connaughton, I. Kolokolov, and V. Lebedev, *Phys. Rev. Lett.* **99**, 084501 (2007).
- [60] J. V. Shebalin, *J. Plasma Phys.* **72**, 507 (2006).
- [61] J. E. Stawarz, C. W. Smith, B. J. Vasquez, M. A. Forman, and B. T. MacBride, *Astrophys. J.* **697**, 1119 (2009).
- [62] H. Politano and A. Pouquet, *Geophys. Res. Lett.* **25**, 273 (1998).
- [63] G. G. Howes and E. Quataert, *Astrophys. J. Lett.* **709**, L49 (2010).
- [64] D. Shaikh and G. P. Zank, *Phys. Plasmas* **12**, 122310 (2005).
- [65] Q. Chen, S. Chen, and G. L. Eyink, *Phys. Fluids* **15**, 361 (2003).
- [66] S. Servidio, W. H. Matthaeus, and V. Carbone, *Phys. Rev. E* **78**, 046302 (2008).