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# Pre-hibernation performances of the OSIRIS cameras onboard the Rosetta spacecraft

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## ABSTRACT

**Context.** The ESA cometary mission Rosetta was launched in 2004. In the past years and until the spacecraft hibernation in June 2011, the two cameras of the OSIRIS imaging system (Narrow Angle and Wide Angle Camera, NAC and WAC) observed many different sources. On 20 January 2014 the spacecraft successfully exited hibernation to start observing the primary scientific target of the mission, comet 67P/Churyumov-Gerasimenko.

**Aims.** A study of the past performances of the cameras is now mandatory to be able to determine whether the system has been stable through the time and to derive, if necessary, additional analysis methods for the future precise calibration of the cometary data.

**Methods.** The instrumental responses and filter passbands were used to estimate the efficiency of the system. A comparison with acquired images of specific calibration stars was made, and a refined photometric calibration was computed, both for the absolute flux and for the reflectivity of small bodies of the solar system.

**Results.** We found a stability of the instrumental performances within  $\pm 1.5\%$  from 2007 to 2010, with no evidence of an aging effect on the optics or detectors. The efficiency of the instrumentation is found to be as expected in the visible range, but lower than expected in the UV and IR range. A photometric calibration implementation was discussed for the two cameras.

**Conclusions.** The calibration derived from pre-hibernation phases of the mission will be checked as soon as possible after the awakening of OSIRIS and will be continuously monitored until the end of the mission in December 2015. A list of additional calibration sources has been determined that are to be observed during the forthcoming phases of the mission to ensure a better coverage across the wavelength range of the cameras and to study the possible dust contamination of the optics.

**Key words.** instrumentation: detectors – space vehicles: instruments – techniques: photometric

## 1. Introduction

The ESA Rosetta mission was launched on 2 March 2004, and since then, the Optical, Spectroscopic, and Infrared Remote Imaging System (OSIRIS, [Keller et al. 2007](#)), which is a two-camera system, performed many different scientific observations. It acquired data during a swing-by with Mars and Phobos in 2007, a fly-by with asteroid (2867) Steins in 2008, two swing-bys with Earth in 2007 and 2009, and a fly-by with asteroid (21) Lutetia in 2010, to mention only the main activities. On 8 June 2011, the spacecraft entered into deep-space hibernation mode. It successfully exited hibernation on 20 January 2014 for the cruise to the rendez-vous with the primary scientific target of the mission: the Jupiter-family comet 67P/Churyumov-Gerasimenko.

Establishing the stability of the instrument through time is vital to performing a correct absolute flux calibration of the cometary data.

In Sect. 2 we therefore study the past performances of the two cameras, the Narrow Angle and Wide Angle Camera (NAC and WAC), through the observations of calibration stars. In particular, we investigate the past responses of calibrator signal through time with different filters to be able to evaluate possible aging effects on the optics and to determine whether the photometric requirements set by the original mission science goals have remained stable. We also inspect the dynamical range of the two CCDs by checking the linear dependence between the stellar signal and the exposure time of observations. The transmission and reflectivity curves of each component of the two cameras, measured during on-ground calibration tests, is used to estimate the efficiency of the instrumentation. We thus check and compare the expected calibrator signals that are deduced from these instrumental efficiency curves and from the observed one, which is computed from images acquired in flight.

Following this investigation, we present in Sect. 3 the robust formulation for the absolute flux calibration of OSIRIS images.

**Table 1.** Basic parameters of the NAC and WAC system and CCD specification (Keller et al. 2007).

	NAC	WAC
Optical design	3-mirror off-axis	2-mirror off-axis
Detector type	2k × 2k CCD	2k × 2k CCD
Angular resolution ( $\mu\text{rad px}^{-1}$ )	18.6	101
Focal length (mm)	717.4	140 (sag)/131 (tan)
Field of view ( $^\circ$ )	$2.20 \times 2.22$	$11.35 \times 12.11$
Telescope aperture	$6.49 \times 10^{-3} \text{ m}^2$	$4.91 \times 10^{-4} \text{ m}^2$
F-number	8	5.6
Spatial scale from 1km (cm $\text{px}^{-1}$ )	1.86	10.1
NAC and WAC CCD		
Pixel size	$13.5 \times 13.5 \mu\text{m}^2$	
Full well	$>120\,000 \text{ e}^- \text{ px}^{-1}$	
System gain	$\approx 3.1 \text{ e}^-/\text{DU}$	
Readout noise (CCD)	$\approx 15 \text{ e}^- \text{ rms}$	
Dark charge generation	$<0.1 \text{ e}^- \text{ s}^{-1} \text{ px}^{-1}$ @ 180 K	
	$\approx 400 \text{ e}^- \text{ s}^{-1} \text{ px}^{-1}$ @ 293 K – (with dithering)	
Readout rate	$1.3 \text{ Mpx s}^{-1}$ ; $650 \text{ kpx s}^{-1}$ per channel	
Readout time (full frame)	$3.4 \text{ s}$ (2 channels)	

**Table 2.** NAC and WAC filter names and wheel positions.

Position	NAC		WAC	
	Wheel 1	Wheel 2	Wheel 1	Wheel 2
1	FFP_UV	FFP_IR	empty	empty
2	FFP_Vis	Orange	Green	Red
3	NFP_Vis	Green	UV 245	UV 375
4	Near-IR	Blue	CS	CN
5	Ortho	Far-UV	UV 295	NH2
6	Fe2O3	Near-UV	OH	Na
7	IR	Hydra	UV 325	OI
8	Neutral	Red	NH	Vis 610

**Notes.** Position of the two filter wheels of all filters for NAC and WAC. In principle, every combination is possible, but only some of them are useful for scientific purposes.

Section 4 is dedicated to the description of the conversion between the signal observed on the surface of solar system bodies and of their correct geometric reflectance.

The NAC is a three-mirror anastigmat system, while the WAC has an off-axis optical configuration obtained with two aspherical mirrors (Keller et al. 2007). In Table 1 we report the main characteristics of the NAC and WAC system, together with the specifications of the CCDs.

A plane-parallel antireflection coated plate, referred to as antiradiation plate (ARP), was added in front of the CCD for radiation shielding.

Both cameras are equipped with two filter wheels placed in front of the CCD. A passband is defined by a combination of two filters, hereafter  $(x, y)$ , where  $x$  is the filter position on the first wheel and  $y$  is the position on the second wheel. In Table 2 we list all filters available for the two cameras by specifying their position on the wheels.

The WAC filters are located in seven of the eight positions of the two filter wheels, from 2 to 8, leaving the first position in both wheels empty. All seven filters of the first wheel can be used by setting the empty position in the second one, and vice-versa.

In the NAC the filter wheels have no empty position. Neutral or clear filters are used in combination with band-pass filters to optimize the camera performances either in the UV, in the visible, or in the IR range. A different focus range for near focus

plate – visible (NFP-Vis) with respect to far focus plate – visible (FFP-Vis) is also planned, which will allow optimizing the resolution of the NAC for observations of the comet surface at a distance shorter than 4 km.

We can therefore define the filters as twins, which can be obtained by combining different clear filters in addition to the same band-pass filter. These combinations share the same passband, but have different efficiencies in general because clear filters have different efficiencies. Together with defining twin combinations, it is also possible to define a baseline filter combination by selecting a representative of the group and deduce the flux calibration of the whole group according to the baseline calibration.

In Appendix A we list and define the physical parameters and units we used.

## 2. Performance of OSIRIS cameras

### 2.1. Expected count rate

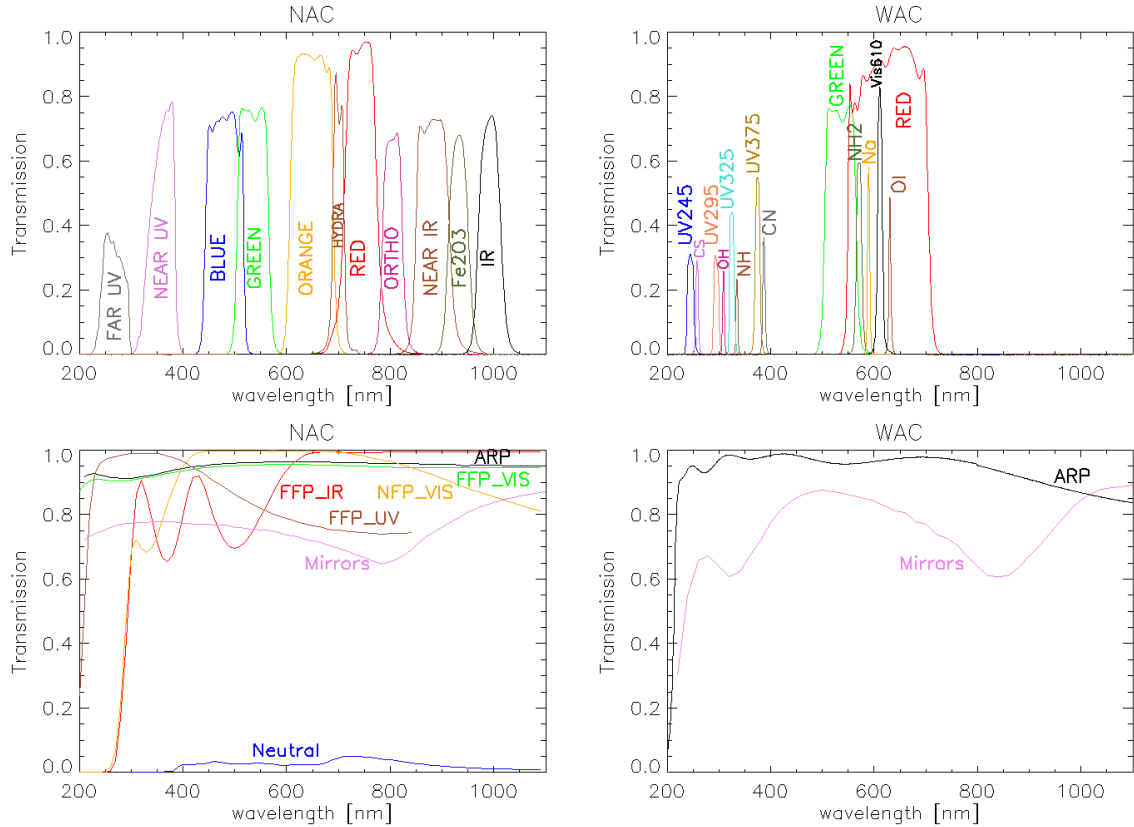
The expected total count rate,  $K_{x,y}$  in  $\text{DN s}^{-1}$ , of a given source on an image acquired through the  $(x, y)$  filter combination is expressed by

$$K_{x,y} = \int_0^\infty \varphi(\lambda) \cdot M^n(\lambda) \cdot F_{x,y}(\lambda) \cdot T_{\text{ARP}}(\lambda) \times Q(\lambda) \cdot \frac{\lambda}{hc} \cdot \frac{1}{\text{IG}} \cdot \text{AP} d\lambda, \quad (1)$$

where

- $\varphi(\lambda)$  is the spectral irradiance of the source at the camera entrance aperture, expressed in  $\text{W m}^{-2} \text{ nm}^{-1}$ ;
- $Q(\lambda)$  is the CCD quantum efficiency expressed in  $\text{e}^-/\text{ph}$ ;
- IG is the inverse gain, which hereafter is taken to be  $3.1 \text{ e}^-/\text{DN}$ ;
- $M^n(\lambda), F_{x,y}(\lambda), T_{\text{ARP}}(\lambda)$  are the mirror spectral reflectance (through  $n$  mirrors:  $n = 2$  for WAC or  $n = 3$  for NAC), filters, and antiradiation plate (ARP) transmissions as a function of the wavelength;
- AP is the telescope aperture:  $4.91 \times 10^{-4} \text{ m}^2$  for the WAC and  $6.49 \times 10^{-3} \text{ m}^2$  for the NAC.

In Fig. 1 we report the reflectivity of the three-mirror (NAC) and two-mirror (WAC) systems together with filters and ARP transmission curves.



**Fig. 1.** Efficiency curves of NAC and WAC optics: transmission curves of clear and band-pass filters, antiradiation plates (ARP), and reflectivity of mirrors.

The integrand of Eq. (1) is the wavelength dependance of the count rate detected by the instrument. In particular,  $Q(\lambda) \times \lambda/hc$  is the way to express the CCD efficiency in  $e^-/(W s)$ , where  $\lambda/(hc) = \lambda/(1.986 \times 10^{-16})$  if  $\lambda$  is expressed in nm.  $Q(\lambda) \times \lambda/(hc \cdot IG)$  hence gives the CCD efficiency in  $DN/(W s)$ .

If the source flux is constant with time, then  $\varphi(\lambda)$  is time independent. Since in principle all instrument optics responses depend on age, the integrand of Eq. (1) and hence  $K_{x,y}$  are therefore time (age) dependent.

## 2.2. Observed count rate

Equation (1) gives the total count rate that is expected on the CCD in each  $(x, y)$  passband of a source with a known spectral irradiance  $\varphi(\lambda)$ . As a first check of the OSIRIS performances, we compared these theoretical values with the real values observed by the instrument.

Two stars were regularly observed by OSIRIS during the past mission phases for calibration purposes: Vega and the double star 16 Cyg A+B. 16 Cyg B is a solar analog and is suitable for measuring the reflectance of small bodies of the solar system. The corresponding spectral irradiances that were taken into account to provide the theoretical values we compared our values with were taken from Bohlin & Gilliland (2004) for Vega and from Burlov-Vasiljev et al. (1995, 1998) for the solar spectrum properly scaled to match 16 Cyg A and B.

### 2.2.1. Aperture photometry

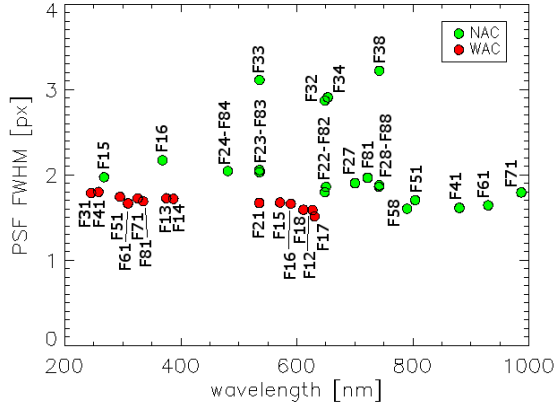
We derived the observed count rate from the calibration stars with the aperture photometry method.

The raw data were processed to remove the bias level. Flat-field correction was applied by means of images that were observed on the ground of the integrating sphere, which were acquired when the cameras were characterized in the laboratory. In-flight flat-field images are regularly acquired by illuminating the closed telescope front doors with internal lamps. However, internal lamps do not guarantee a completely flat source, and this method is only useful to monitor the flat-field stability, but not to create new reference images. Since no variation was detected on in-flight images, laboratory images are still used for the flat-field correction.

In the WAC images the 16 Cyg A+B pair is not separated, which forced us to apply the method to both stars at the same time. In the NAC images the 16 Cyg A and B centroids are separated by about ten pixels. This means that we can measure fluxes of 16 Cyg B alone only when the aperture radius is verified to be smaller than 5 pixels, to avoid contamination from 16 Cyg A.

On the CCD, different filters show different typical point spread function (PSF) profiles (see for example Figs. 2 and 3), so that the best aperture radius to use for the aperture photometry depends on the filter, which in many cases is larger than 5 pixels. For consistency between different cameras and among different filters, we therefore decided to derive the photometry from both stars in the NAC in the same way as we have to do with WAC images. See Appendix B for more details on the double star 16 Cyg A+B.

We finally had to define a general criterion to find the best value for the aperture radius. By investigating the stellar integral profile obtained at different aperture radius, we typically expect to find a plateau for aperture radii larger than a critical value after the background level is removed. For OSIRIS images, a typical aperture growth curve of Vega is reported in Fig. 4, where



**Fig. 2.** PSF FWHM as number of pixels (geometric average of  $x$  and  $y$  direction values) in OSIRIS filters deduced from Vega observations. The largest size of four NAC filters is acquired by combining a near focus plate that can optimize the performances of the NAC camera for a target that is at a distance smaller than 4 km.

the sky level is estimated in a region here defined by sky apertures between 50 and 60 px. For both NAC and WAC, the sky background is normally found to be within  $\pm 3$  DN, with a standard deviation of about 8.5 DN on NAC and 6.7 DN on the WAC, probably due to coherent noise. The plateau indicates that the total flux of the star was taken into account.

In some OSIRIS images this cannot be achieved. We often found an oblique (instead of a horizontal) asymptote, regardless of our definition of the sky annuli. We ascribed this to an on-axis contamination, maybe due to diffuse light by optics micro-roughness or to ghosts.

The NAC ghosts are well described in [Dohlen et al. \(2010\)](#). First-generation ghosts are caused by the light reflection between pairs of optical surfaces. The most intense on-axis ghost, for example, can be produced by the reflection of the incident light on the surface of the CCD, reflected again by the nearest surface of the ARP. In the [800–1000] nm IR range the NAC CCD reflectivity increases from 10 to 30%. The ghost relative intensity, with respect to the scientific image, is thus three times higher at 1000 nm than at 800 nm. The two filters most affected by the oblique asymptote are the reddest ones, *F61* and *F71*, which are centered at about 930 and 990 nm.

However, the CCD reflectivity and the ARP reflectivity in the whole [250–400] nm range are higher than at 1000 nm, which means that the ghost problem should affect the UV range much more than the IR, while we find that the contamination in UV NAC filter, where present, is not so evident. Therefore, further investigation is needed on this topic.

In the meanwhile, we had to find a new definition for the aperture radius for the NAC-*F61* and NAC-*F71* filters: we selected the smallest aperture radius  $r$  such that the flux in the annulus  $[r-1, r]$  is below half the sky standard deviation. After that we considered the remnant contribution as background even if with a non-zero average.

### 2.2.2. CCD linearity and stability

We investigated the stability of the results of the aperture photometry method during different observational phases of the mission (from the Mars swing-by in 2007 to the Lutetia fly-by in 2010).

In Table 3 we report the mean value of the count rate computed from Vega images with different filter combinations, the maximum deviation, and the typical exposure times.

For the two NAC (*F16* FFP-UV\_Near-UV, and *F82* Neutral\_Orange) and two WAC filters (*F12* Red, and *F18* Vis610) the signal linearity with increasing exposure time was also investigated. The shortest exposure time is imposed by the shutter response tolerance, the longest time is defined by the CCD saturation level.

The OSIRIS CCDs are equipped with an antiblooming system, designed to “extract” overabundant electrons from the pixel potential well. This system is needed for the cometary target to allow overexposed images of the nucleus without saturation artifacts. This is necessary to acquire high signal-to-noise (S/N) data of the dust and gas in the inner coma against the brilliant nucleus. The saturation level was required to be about  $120\,000\text{ e}^-/\text{px} \approx 40\,000\text{ DN/px}$  (see [Keller et al. 2007](#); [Thomas et al. 1998](#)) so that the pixel linearity is guaranteed only below this limit. Above the saturation, the CCD cannot collect any other electron, and we will find a constant signal regardless of the exposure time.

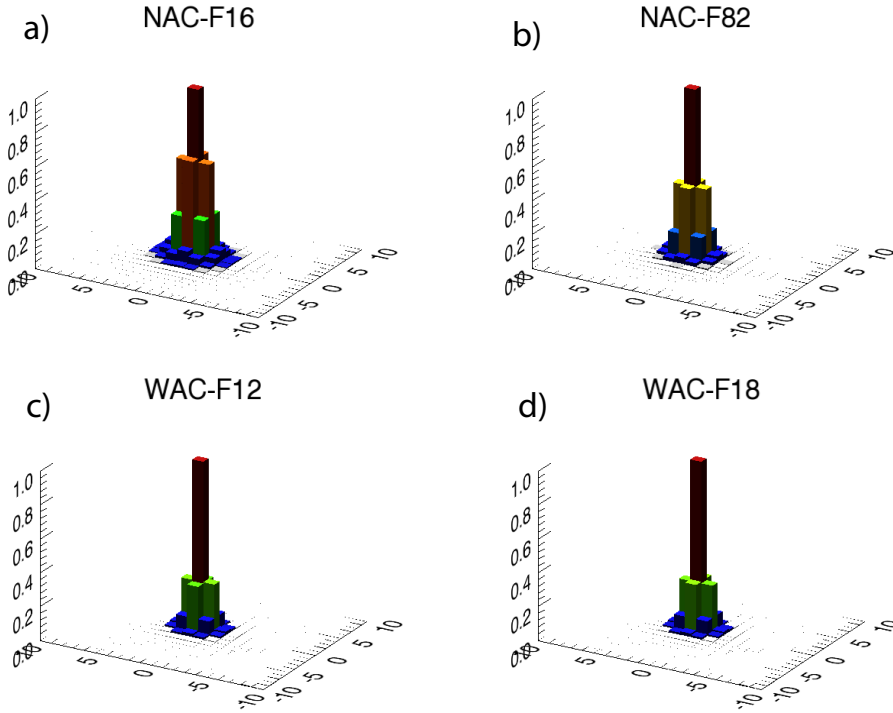
For point sources, which are dominated by the optics PSF, the peak value of the source must not exceed the saturation value, while the total flux of the source is split into more pixels and the total number of electrons collected can obviously exceed the limit. When the peak signal reaches the saturation levels, the antiblooming system prevents overabundant charges of only the overexposed pixels to spill over into the neighborhood above a given exposure time. Nearby unsaturated pixels, however, can continue to collect electrons, so that we would still experience an increase of the total point source signal with exposure time, even with a lower rate.

To study the linearity of the CCD response, we defined as outliers and consequently excluded from the data set the count rates that deviate from the mean value by more than three times the standard deviation. The iteration of this method defined a time interval where the linearity is satisfied. Figure 5 shows the distribution of Vega observations (reported as both counts and count rate) vs. exposure time in this interval. In Table 3 the time interval is specified for these four filters.

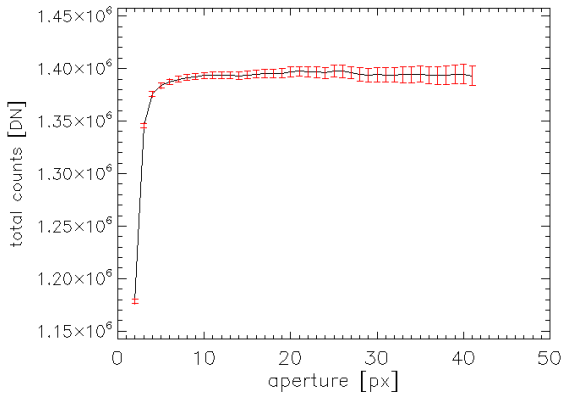
Table 3 shows that almost all the filters are stable within  $\pm 1.5\%$ . In only one case, WAC-*F51* at 295 nm, there is a deviation of  $+2.79\%$ , while in only two other filters, NAC-*F41* at 880 nm and WAC-*F61* at 308.5 nm, deviations exceed  $1.5\%$ . Moreover, we also verified that the data considered do not show any evident aging effect since there is no systematic trend with date: the data dispersion of about  $\pm 1.5\%$  is found both in case of stellar observations acquired during the same day and during the whole OSIRIS activity period. This dispersion is also consistent with the photometry precision requirement of about 2% set by the original mission science goals ([Keller et al. 2007](#)).

From this investigation we found that the peak level of counts on a single pixel in raw images is above 57 000 DN for both NAC filters investigated, while it is about 47 000 and 50 000 DN for the two WAC filters investigated, well above the 40 000 DN limit foreseen. We verified for WAC that the centroids of the PSF were located in different positions on the CCD, with different efficiencies: if the saturation limit was found to be at 47 000 DN, the star always fell on some few cold pixels. This difference was then corrected by flat fielding. The difference between NAC and WAC values suggest that the two CCDs behave differently. For NAC-*F82* we did not exclude any observation in the data-set with the longest exposure times, so that we did not confirm the saturation limit in this case. For new tests that are scheduled to





**Fig. 3.** Shape of the Moffat fit of the PSF of OSIRIS camera filters in the following combinations: **a)** NAC-*F16* (FFP-UV\_Near-UV); **b)** NAC-*F82* (Neutral\_Orange); **c)** WAC-*F12* (Red); and **d)** WAC-*F18* (Vis610).



**Fig. 4.** Typical curve of the aperture flux growth. This curve was specifically obtained from Vega observations with the NAC *F16* filter.

study the CCD linearity, a wider range in  $t_{\text{exp}}$  for NAC-*F82* must be considered.

In Fig. 3 we show the PSF for two NAC and two WAC filters in the case where the centroid of the PSF lies exactly on the center of a pixel. For the four selected filters the ratios between the peak value and the total integrated signal are about 17.5% for NAC-*F16*, 24.5% for NAC-*F82*, 34.5% for WAC-*F12*, and 34.3% for WAC-*F18* of the total signal.

### 2.3. Count rate: observed vs. expected

In Fig. 6 we show the ratio between observed and expected counts in each filter for Vega and 16 Cyg A+B. The count rate obtained for Vega and 16 Cyg A+B observations acquired during all mission phases was compared with the expected rate that we computed from the Vega reference spectrum (Bohlin & Gilliland 2004, CALSPEC version 004) and solar reference spectra (Burlov-Vasiljev et al. 1995, 1998), respectively, by means of Eq. (1). For this last source we computed a conversion factor of between 16 Cyg A+B and solar fluxes, based on the

different visual absolute magnitudes, to allow a direct comparison with the solar spectrum. See Appendix B for further details.

The error bars in this plot were computed by propagating the photon noise deduced on the original images with errors defined on each instrument element. Horizontal bars refer to the FWHM of each passband.

The total relative (percentage) instrumental error is dominated by the CCD QE uncertainty.

Several considerations arise from the analysis of the plot:

- In the UV range, that is, below 400 nm, fewer counts are observed than expected.
- In the visible range, the WAC filters counts agree with the predicted counts. The NAC filter counts, however, seem to exceed the expected counts (in some cases they are even higher than the vertical error bars).
- In the near-IR range, above 800 nm, fewer counts are observed than predicted.
- The NAC twin filter values agree within error bars. However, we found some systematic differences among NAC FFP-Vis, NAC NFP-Vis, and NAC Neutral (second, third, and eighth position on filter wheel 1, respectively).
- The agreement between observed and expected counts is verified from both Vega and 16 Cyg A+B counts in the range 470–890 nm.
- The agreement between Vega and 16 Cyg A+B data is good above 470 nm.
- The scatter among filters and between different stars is larger in the UV range. 16 Cyg A+B data are systematically lower than Vega data.
- Around WAC *F71* (centered on 325 nm) the plot behaves strangely in both Vega and 16 Cyg A+B data. This scatter is not due to an anomalous observation because the photometric values were computed for this filter, like for others, as the average of many observation in different phases, giving a positive and a negative maximum percentage error of +1.11% and –1.36%, respectively.

**Table 3.** OSIRIS NAC and WAC filter characteristics and Vega observations.

Filter	Central $\lambda$ nm	FWHM nm	$T_{\text{exp}}$ s	Mean count rate DN/s	+err %	-err %	Observations	Aperture px
NAC-F15	267.5	53	0.400	2.571e+005	1.07	0.95	6	6–8
NAC-F16	368.5	33	0.020–0.256	1.415e+006	0.59	0.75	30	8–10*
NAC-F22	648.5	83	0.020	7.172e+006	1.19	0.93	5	9–11
NAC-F23	535.5	61	0.020	7.265e+006	0.80	0.69	5	10–12
NAC-F24	481.0	72	0.020	9.369e+006	0.15	0.30	5	10–11
NAC-F27	700.5	21	0.080	1.366e+006	1.29	0.91	5	9–10
NAC-F28	742.0	62	0.025	3.447e+006	0.27	0.22	5	9–11
NAC-F32	648.5	83	0.040	7.496e+006	0.41	0.83	5	13–14
NAC-F33	535.5	61	0.040	7.581e+006	1.38	1.25	5	13–15
NAC-F38	742.0	62	0.060	3.565e+006	1.10	0.61	4	13–14
NAC-F41	880.0	62	0.100	9.906e+005	1.76	1.03	5	10–11
NAC-F51	804.5	39	0.100	9.895e+005	1.16	0.79	5	9–10
NAC-F58	790.5	23	1.000	7.631e+004	0.50	0.34	6	9–10
NAC-F61	929.5	39	0.350	3.223e+005	0.71	1.15	5	13–15
NAC-F71	987.0	38	0.800	1.092e+005	1.43	1.50	5	16–17
NAC-F82	650.5	81	0.070–1.297	1.941e+005	0.58	0.76	32	9–11*
NAC-F83	535.5	59	0.500	2.033e+005	1.32	0.80	5	8–9
NAC-F88	742.0	60	0.500	1.752e+005	1.37	1.43	5	9–10
WAC-F12	627.0	154	0.010–0.150	1.169e+006	1.39	1.35	66	8–11*
WAC-F13	375.0	8	6.000	1.875e+004	1.23	0.85	8	11–12
WAC-F14	387.0	4	6.000	1.588e+004	1.21	0.93	9	10–12
WAC-F15	571.0	10	1.500	8.142e+004	0.73	0.53	9	9–11
WAC-F16	589.5	3	3.500	3.030e+004	0.77	0.55	8	9–10
WAC-F17	630.5	3	3.500	2.029e+004	0.85	0.97	9	7–9
WAC-F18	611.5	9	0.233–1.900	8.288e+004	0.52	1.14	29	8–10*
WAC-F21	535.5	61	0.200	5.910e+005	0.68	0.28	9	9–11
WAC-F31	245.5	13	15.000	5.308e+003	0.85	1.27	9	12–13
WAC-F41	258.0	4	50.000	1.749e+003	1.05	0.84	8	12–14
WAC-F51	295.0	10	25.000	3.060e+003	2.79	1.79	6	12–13
WAC-F61	308.5	3	60.000	1.273e+003	1.77	1.85	8	11–13
WAC-F71	325.5	9	20.000	4.376e+003	1.11	1.36	9	12–13
WAC-F81	335.0	4	80.000	1.422e+003	0.86	0.69	8	11–12

**Notes.** This table refers to the observations of Vega during different mission phases, from the Mars swing-by in 2007 to the Lutetia fly-by in 2010. For each filter we report the central wavelength of the whole instrumental and filter passband in nm; the FWHM of each instrumental and filter passband in nm; the exposure time needed for the Vega observation, and for the four investigated filters, the linearity interval in seconds; the average count rate in DN s<sup>-1</sup> for Vega observations; the positive and negative maximum deviation (expressed in %) with respect to the mean value; the total number of observations; the typical values of the apertures among all observations; (\*) for each of the four additional filters, this refers only to observations performed with a single particular exposure time within the interval specified in the fourth column.

16 Cyg A+B data can be affected by an incorrect computation of the conversion from solar fluxes (computed from  $10^{-0.4(V_{\text{SA}} - V_{\odot})}$ ). However, the match of 16 Cyg A+B data with Vega data above 470 nm suggests that this assumption is good at least in the visible and near-infrared range of the spectrum. In the UV range, there may be also some differences due to the imperfect match between 16 Cyg A and 16 Cyg B spectra and the solar spectrum, which can be variable in these wavelengths by up to 20% (Colina et al. 1996). Moreover, below 325 nm, 16 Cyg A+B counts are maybe unreliable as a result of the decreasing signal in this wavelength range that is typical of solar analog stars, which means that we can only trust the Vega counts.

Variability among twin filters that share the same passband (i.e., NAC F82-F22-F32, NAC F83-F23-F33, NAC F88-F28-F38, etc.) are probably due to different shapes of the actual transmission curves of clear filters (NAC FFP-Vis, NAC NFP-Vis, and NAC Neutral) with respect to laboratory data of the same filters obtained on the ground before the launch.

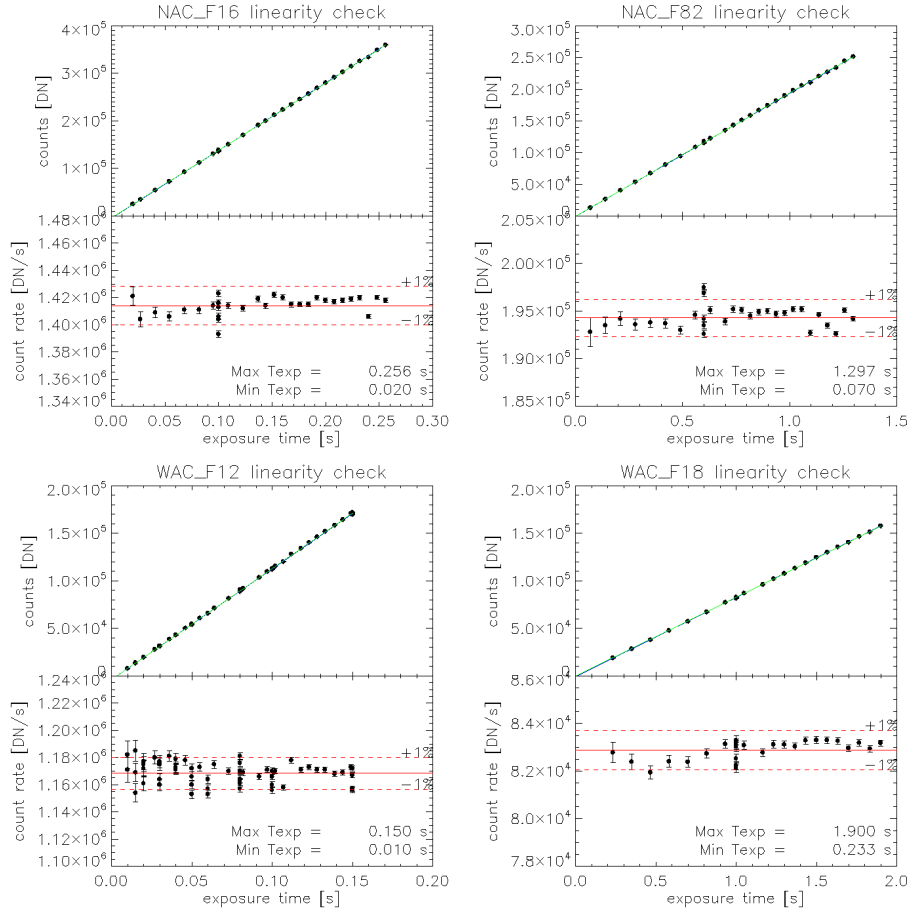
The lower count rate in the near IR can in principle be due to an incorrect computation of the counts within the selected aperture. For many NAC filters the aperture radii selected for the computation of Vega flux through the aperture photometry

method were found to be 7–9 pixels. We recall here that for these filters we always found a plateau. However, throughout all the past mission phases, the signal was always found to increase for the NAC IR filters F61 and F71; it lay on an oblique asymptote. Taking into account a possible on-axis contamination, we decided to stop at an aperture of about 10–12 pixels for F61 and at 15–17 pixels for F71, depending on the background noise level on the different images. For the sake of completeness, we investigated wider apertures for the NAC-F61 and F71 filters. We found that even 40-pixel apertures are not enough to obtain the expected counts, hence wider apertures cannot improve the calibration in either case.

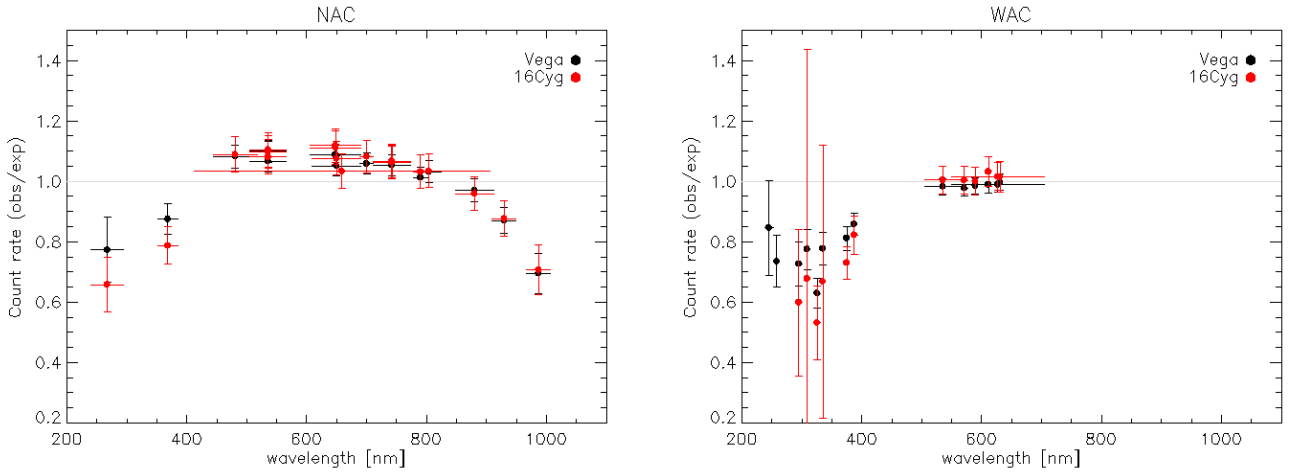
### 3. Absolute flux calibration

We call the factor needed to convert observed count rates per pixel into physical units the abscal factor (i.e., spectral radiance, expressed in W m<sup>-2</sup> nm<sup>-1</sup> sr<sup>-1</sup>) to obtain the absolute instrument calibration.

We present photometric calibration factors for each passband. We used images of stellar targets such as Vega and the pair 16 Cyg A+B. According to each filter combination ( $x, y$ ),



**Fig. 5.** Linearity measurements obtained by the Vega observations in terms of DN or DN/s vs exposure time. For each filter, the *upper panels* show counts vs. exposure time in the particular filter, the *lower panels* show the count rate vs. the exposure time. The investigation was made with the following filters: **a)** NAC-F16 (FFP-UV\_Near-UV); **b)** NAC-F82 (Neutral\_Orange); **c)** WAC-F12 (Red); and **d)** WAC-F18 (Vis610).



**Fig. 6.** Ratio between observed and expected counts from the CCD with a particular filter setting. Vertical error bars refer to photon noise and the propagation of instrumental uncertainties. Horizontal bars refer to the FWHM of each passband. Red data refer to counts computed from 16 Cyg A+B images and solar spectrum (Burlov-Vasiljev et al. 1995, 1998) scaled to 16 Cyg A+B magnitude; black data refer to Vega images and Vega spectrum (Bohlin & Gilliland 2004).

we can measure the total count rate due to the star (or pair),  $K_{x,y}$ , from these images.

The spectral radiance that we expect on each pixel with the same filter combination ( $x, y$ ) is the weighted average of the incoming spectral radiance of the target star, with the instrumental response as the weight, divided by the pixel size (pxsz) of

the CCD expressed in steradians ( $3.547\text{e-}010$  sr for NAC and  $9.982\text{e-}009$  sr for WAC):

$$\phi_{x,y} = \frac{\int_0^{+\infty} \varphi(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \frac{\lambda}{hc} \frac{\text{AP}}{\text{IG}} d\lambda}{\text{pxsz} \int_0^{+\infty} M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \frac{\lambda}{hc} \frac{\text{AP}}{\text{IG}} d\lambda}.$$



**Table 4.** Absolute flux calibration of the OSIRIS NAC and WAC filters.

Filter	Central $\lambda$ nm	FWHM nm	Abscal factor $\text{DN s}^{-1} \cdot (\text{Wm}^{-2} \text{nm}^{-1} \text{sr}^{-1})^{-1}$	+err %	-err %	Abscal reflectance factor $\text{DN s}^{-1}$	+err %	-err %
NAC F15	267.5	53.0	2.428e+006	16.7	16.7	1.434e+005	15.3	15.3
NAC F16	368.5	33.0	1.308e+007	5.6	5.7	4.566e+006	5.6	5.7
NAC F21	659.0	492.0	6.190e+008	3.3	3.2	2.830e+008	1.7	1.7
NAC F22	648.5	83.0	1.182e+008	2.5	2.4	5.892e+007	2.5	2.4
NAC F23	535.5	61.0	6.677e+007	2.4	2.5	4.01e+007	2.4	2.5
NAC F24	481.0	72.0	6.421e+007	2.7	2.5	4.086e+007	2.7	2.5
NAC F26	368.5	33.0	1.232e+007	5.6	5.7	4.307e+006	5.6	5.7
NAC F27	700.5	21.0	2.824e+007	2.4	2.4	1.270e+007	2.4	2.4
NAC F28	742.0	62.0	8.544e+007	2.3	2.3	3.495e+007	2.3	2.3
NAC F31	653.5	479.0	6.297e+008	3.2	3.2	2.907e+008	3.9	3.9
NAC F32	648.5	83.0	1.235e+008	2.5	2.4	6.156e+007	2.5	2.4
NAC F33	535.5	61.0	6.959e+007	2.5	2.6	4.169e+007	2.5	2.6
NAC F34	481.0	72.0	6.741e+007	2.6	2.5	4.290e+007	2.6	2.5
NAC F35	289.5	13.0	3.157e+005	14.9	14.9	3.980e+004	14.5	14.5
NAC F36	371.5	27.0	1.071e+007	5.6	5.7	3.765e+006	5.5	5.6
NAC F37	700.5	21.0	2.942e+007	2.4	2.4	1.323e+007	2.4	2.4
NAC F38	742.0	62.0	8.842e+007	2.3	2.3	3.618e+007	2.3	2.3
NAC F41	880.0	62.0	4.024e+007	3.4	3.2	1.209e+007	3.4	3.2
NAC F51	804.5	39.0	3.159e+007	2.7	2.5	1.141e+007	2.7	2.5
NAC F58	790.5	23.0	2.326e+006	2.9	2.5	8.626e+005	2.9	2.5
NAC F61	929.5	39.0	1.444e+007	4.3	4.3	3.894e+006	4.3	4.3
NAC F71	987.0	38.0	5.763e+006	10.8	10.7	1.384e+006	10.7	10.6
NAC F81	722.0	206.0	1.937e+007	3.1	3.1	8.938e+006	3.8	3.8
NAC F82	650.5	81.0	3.246e+006	2.7	2.4	1.607e+006	2.7	2.4
NAC F83	535.5	59.0	1.869e+006	2.5	2.4	1.120e+006	2.5	2.4
NAC F84	481.0	72.0	1.991e+006	2.7	2.5	1.270e+006	2.7	2.5
NAC F86	378.0	18.0	4.980e+004	5.1	5.2	1.819e+004	5.0	5.1
NAC F87	700.5	21.0	1.313e+006	2.4	2.4	5.891e+005	2.4	2.4
NAC F88	742.0	60.0	4.286e+006	2.7	2.8	1.756e+006	2.7	2.8
WAC F12	627.0	154.0	4.810e+008	2.5	2.6	2.52e+008	2.5	2.6
WAC F13	375.0	8.0	4.684e+006	4.8	4.8	1.774e+006	4.8	4.8
WAC F14	387.0	4.0	2.538e+006	3.8	3.9	8.826e+005	3.8	3.8
WAC F15	571.0	10.0	2.557e+007	2.3	2.1	1.494e+007	2.3	2.1
WAC F16	589.5	3.0	1.045e+007	2.1	2.1	5.823e+006	2.1	2.1
WAC F17	630.5	3.0	8.594e+006	2.1	2.3	4.435e+006	2.1	2.3
WAC F18	611.5	9.0	3.196e+007	2.1	2.2	1.724e+007	2.1	2.2
WAC F21	535.5	61.0	1.523e+008	2.3	2.1	9.131e+007	2.3	2.1
WAC F31	245.5	13.0	1.370e+006	23.6	23.5	3.449e+004	20.9	20.7
WAC F41	258.0	4.0	4.686e+005	13.9	13.8	1.952e+004	12.9	12.8
WAC F51	295.0	10.0	8.483e+005	11.8	11.7	1.443e+005	11.7	11.6
WAC F61	308.5	3.0	3.592e+005	9.8	9.7	7.412e+004	9.6	9.5
WAC F71	325.5	9.0	1.312e+006	8.6	8.6	3.690e+005	8.6	8.5
WAC F81	335.0	4.0	4.447e+005	7.7	7.6	1.357e+005	7.7	7.6

**Notes.** In this table we report the absolute flux and reflectance calibration of all OSIRIS NAC and WAC filter combinations, obtained using the whole data set of Vega and 16 Cyg A+B observations during the different mission phases from the Mars swing-by in 2007 to the Lutetia fly-by in 2010. For each filter we report the central wavelength of the whole instrumental and filter passband in nm; the FWHM of each instrumental and filter passband in nm; the abscal factor needed to convert the image count rate into spectral radiance; its positive and negative percentage error; the abscal reflectance factor needed to convert the image count rate into reflectance; its positive and negative percentage error.

Consequently, the abscal factor,  $A_{x,y}$ , which links the total count rate ( $K_{x,y}$ ) on the CCD for a particular filter combination to the spectral radiance within the passbands,  $\phi_{x,y}$ , expressed in  $\text{W m}^{-2} \text{nm}^{-1} \text{sr}^{-1}$ , is defined by

$$A_{x,y} = K_{x,y} / \phi_{x,y}. \quad (2)$$

In first approximation  $A_{x,y}$  is filter dependent (and hence in principle time (age) dependent) and star independent.

Because AP,  $hc$ , and IG are constant with wavelength and can be simplified, the abscal factor is

$$A_{x,y} = \frac{K_{x,y} \text{pxsz} \int_0^{+\infty} M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}{\int_0^{+\infty} \varphi(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}. \quad (3)$$

The abscal factor is expressed in  $\text{DN s}^{-1} \cdot (\text{Wm}^{-2} \text{nm}^{-1} \text{sr}^{-1})^{-1}$ .

In Table 4 we report the abscal factor values for all different filters computed from this equation.

$Q(\lambda)$  is roughly constant on the visible passbands and can be simplified from both integrals in Eq. (3); this means in particular that the abscal factors are little affected by the imperfect knowledge of the CCD quantum efficiency in the visible range (approximately [400–1000] nm).

### 3.1. Twin combinations

As said in the Introduction, we can calibrate scientific images acquired with any of the filters of a twin group by using

**Table 5.** List of NAC filters without stellar calibration.

Camera	Filter to calibrate	Baseline filter
NAC	26	16
NAC	31	21
NAC	34	24
NAC	35	15
NAC	36	16
NAC	37	27
NAC	81	21
NAC	84	24
NAC	86	16
NAC	87	27

**Notes.** To calibrate these filters, we used the count rate from the baseline filters listed in the right column.

the baseline calibration of a representative filter of the group. For example, we can calibrate both NAC-*F87* and NAC-*F37*, or NAC-*F84* and NAC-*F34*, simply by using the baseline calibration obtained with filter combinations NAC *F27* or NAC *F24*, respectively. In Table 5 we show all the filters for which this is possible and the selected baseline filter.

In the following we describe the equations used to properly infer the absolute calibration in these cases. As an example, we use the calibration of NAC-*F84* images through the NAC-*F24* observation of the calibration star. The filter in the second position (2) on the NAC first filter wheel (FFP-Vis: far focus plate – visible) has a mean efficiency  $\approx 95\%$ , while filter 8 (neutral) has  $\approx 5\%$ . Accordingly, the two considered filters refer to Neutral\_Blue or FFP-Vis\_Blue (coded as NAC-*F84* or NAC-*F24*), respectively. To compute the absal factor of the filter combination  $A_{84}$ , we cannot use the usual expression

$$A_{84} = \frac{K_{84} \text{pxsz} \int_0^{+\infty} M^n(\lambda) F_{84}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}{\int_0^{+\infty} \varphi(\lambda) M^n(\lambda) F_{84}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}$$

since we did not observe calibration stars through the  $F_{84}$  filter, we do not know the  $K_{84}$  count rate, but only  $K_{24}$ . However, we expect from Eq. (1) that the counts of a star through a filter are proportional to the integral of the stellar spectrum multiplied by the passbands, so that

$$\frac{K_{84}}{K_{24}} = \frac{\int_0^{+\infty} \varphi(\lambda) M^n(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) F_{84}(\lambda) \lambda d\lambda}{\int_0^{+\infty} \varphi(\lambda) M^n(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) F_{24}(\lambda) \lambda d\lambda}. \quad (4)$$

This means that we can express  $A_{84}$  as

$$A_{84} = K_{24} \frac{\text{pxsz} \int_0^{+\infty} M^n(\lambda) F_{84}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}{\int_0^{+\infty} \varphi(\lambda) M^n(\lambda) F_{24}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}.$$

This method was applied to complete the absolute flux calibration presented in Table 4 for filters without a direct observation of the calibration stars.

### 3.2. Theoretical and real absal factor

Since the expected counts from a source should be expressed as in Eq. (1), the theoretic absal factor should also be expressed as

$$A_{x,y}^{\text{theoretic}} = \text{pxsz} \int_0^{+\infty} M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \frac{\lambda}{hc} \frac{\text{AP}}{\text{IG}} d\lambda.$$

In this form we can see that the absal factor should be calibrator independent. However, the expected counts in general may be not equal to the observed counts, as shown in Sect. 2.3. This is why the previous formulation is only the theoretical formulation and needs to be corrected for each variation on the instrumental response with time (age).

In the following, we call  $C_{x,y}$  the ratio between observed and expected counts in the  $(x, y)$  filter. It is defined by

$$C_{x,y} = \frac{K_{x,y}^{\text{observed}}}{\int_0^{+\infty} \varphi(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \frac{\lambda}{hc} \frac{\text{AP}}{\text{IG}} d\lambda}. \quad (5)$$

This equation defines the correction factor that we need to apply to the optics profiles so that Eq. (1) gives the observed count rate. However, with this definition we only assume single correction factors, constant on wavelengths, for each  $(x, y)$  filter combination.

In this approximation the absal factor we should use to convert our images becomes

$$A_{x,y}^{\text{real}} \cong C_{x,y} A_{x,y}^{\text{theoretic}}. \quad (6)$$

We stress that this is not an exact equivalence. We should not only correct for  $C_{x,y}$  factors, constant on wavelength, but we should find a correction profile  $C(\lambda)$ , dependent on the wavelength, deduced by  $C_{x,y}$  data, that can correct the shape of every passband profile.

However, where the response diversity is a constant factor over the passband (or maybe even a slow symmetrical variation over the passband), the theoretical source flux weighted mean is the same as the observed one, and the previous equivalence is exact. This is verified for OSIRIS filters in the [470–900] nm range.

In this case, we can write

$$A_{x,y}^{\text{real}} = K_{x,y}^{\text{observed}} \text{pxsz} \times \frac{\int_0^{+\infty} M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}{\int_0^{+\infty} \varphi(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}. \quad (7)$$

Therefore, we can convert the count rate into the correct physical units in each passband  $(x, y)$  because the absal factor compensates for every variation in the optics efficiencies:

$$\phi_{x,y} = \frac{K_{x,y}^{\text{expected}}}{A_{x,y}^{\text{theoretic}}} \cong \frac{K_{x,y}^{\text{observed}}}{A_{x,y}^{\text{real}}}$$

deduced from Eq. (2).

In contrast, where differences between actual and laboratory measured profiles are asymmetric over the passbands, Eq. (7) is valid only in a first approximation, and the goodness of the approximation depends on the real shape of the passbands. In this case we could also find different absolute flux calibrations for every source observed. This may be the reason why we see a mismatch between Vega and 16 Cyg A+B data in the UV range in Fig. 6. The steep UV behavior of the plot suggests that in this range, NAC and WAC passbands may have profiles with a shape different from that considered for the computation. These different shapes give an incorrect weight for computing the weighted arithmetic means that is at the base of the concept of Eq. (3).

These considerations are of fundamental importance for investigating possible aging or dust contamination of the optics during the last part of the mission around the comet, since we expect a wavelength-dependent variation of optical efficiencies in these cases. Therefore, we will have to continuously check this

calibration: the computation of abscol factors may change with time, and the calibration will be good only because these factors compensate for efficiency variations. Moreover, an accurate computation of the correction profile  $C(\lambda)$  deduced from  $C_{x,y}$  must be implemented to improve the UV calibration.

One way to monitor the goodness of the approximation in Eq. (6) is to compute, for example, the flux calibration by means of different sources, with different spectra  $\varphi(\lambda)$ .

### 3.3. List of UV calibrators

We should therefore observe different stars with different spectral type to be able to check the consistency of the calibration independently of the star spectrum so that we can try to fix calibration issues after the awakening of OSIRIS. As we stated before, 16 Cyg A and B are solar analogs, with a very low S/N in the UV, which is the reason why at the moment only Vega is used at these wavelengths. We need stars different from Vega (A0V) and hotter than a solar analog (G2V).

In Table 6 we present a list of stars in the Bruzual-Persson-Gunn-Stryker (BPGS) atlas with spectral type earlier than A3. These stars are cross-correlated with the TD1 catalog (Thompson et al. 1978) and the corresponding magnitudes are also provided (Hayes & Latham 1975).

The BPGS atlas consists of 175 spectra of stars covering a complete range of spectral types and luminosity classes. Each individual spectrum covers the [229–25 600] Å wavelength region. The fluxes were arbitrarily normalized to a visual magnitude of zero. More details about the spectra can be found in Gunn & Stryker (1983).

## 4. Reflectance calibration

If we are observing a small body in the solar system, at the first approximation and at least in the visible range, we detect that the solar spectrum is diffused from the body surface. Considering OSIRIS data, we can remove the solar contribution on each image by computing the solar flux that falls inside each filter passband to obtain the spectrophotometric reflectance of the observed object inside each passband. We are able, therefore, to investigate its optical properties and give hints on its surface mineralogical composition.

For example, we wish to scale the image count rate so that if the target is observed at zero phase angle, it will give the value of the geometric albedo on the given passband. For that purpose, we recall some photometric definitions.

The geometric albedo is the ratio of a body brightness at zero phase angle to the brightness of a perfectly diffusing (Lambertian) disk with the same position and apparent size as the body. A Lambertian surface follows the cosine emission law (Lambert 1760): the intensity seen by an observer from an ideal diffusing surface is proportional to the cosine of the angle  $\theta$  between the observer's line of sight and the normal to the surface, with peak luminous intensity ( $I_{\text{peak}}$ ) in the normal direction. The incident solar flux is equal to the total diffused flux that can be computed from the peak intensity by integrating the cosine law on the semi-sphere defined by the normal direction, obtaining  $F_{\text{tot}} = \pi I_{\text{peak}}$ .

The geometric albedo in a  $(x, y)$  given passband is expressed by

$$p_{x,y} = \frac{I_{x,y}}{\varphi_{\odot,x,y}(r)/\pi}, \quad (8)$$

where  $I_{x,y}$  is the body intensity at zero phase angle,  $\varphi_{\odot,x,y}(r)$  is the incident solar spectral irradiance at the body distance  $r$  from the Sun, and the factor  $\pi$  comes from the integral on the semi-sphere of the Lambert cosine law.

We can accordingly define the reflectance  $R$  of a solid object illuminated by the Sun at a given phase angle  $\alpha$  as

$$R_{x,y}(\alpha) = \frac{I_{x,y}(\alpha)}{\varphi_{\odot,x,y}(r)/\pi}, \quad (9)$$

where  $I_{x,y}(\alpha)$  is the intensity emitted by the body in the  $(x, y)$  band at the phase angle  $\alpha$  and  $\varphi_{\odot,x,y}(r)$  is the peak of the solar flux diffused by the Lambertian disk with the same cross-section of the object at the distance  $r$  from the Sun in the same band.

From the solar spectrum at 1 AU (Burlov-Vasiljev et al. 1995, 1998) we can compute the expected solar count rate at 1 AU, according to Eq. (1):

$$K_{x,y\odot} = \int_0^{+\infty} \varphi_{\odot}(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \frac{\lambda}{hc} \frac{\text{AP}}{\text{IG}} d\lambda,$$

expressed in DN s<sup>-1</sup>, where  $\varphi_{\odot}(\lambda)$  is the solar spectral irradiance at 1 AU expressed in W m<sup>-2</sup> nm<sup>-1</sup>. The solar spectral irradiance at 1 AU inside a particular filter  $(x, y)$  is expressed by the weighted mean:

$$\varphi_{x,y\odot} = \frac{\int_0^{+\infty} \varphi_{\odot}(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}{\int_0^{+\infty} M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}.$$

After dividing it by  $\pi$  and scaling it to the distance  $r$  of the target to the Sun expressed in astronomic units, we obtain the value  $\varphi_{\odot}(r)/\pi$ , expressed in W m<sup>-2</sup> nm<sup>-1</sup> sr<sup>-1</sup>, used to obtain the (geometric) reflectance within a given passband. Note that the reflectance is a adimensional number.

As already cited, we used the spectrum of the Sun measured by Burlov-Vasiljev et al. (1995, 1998). Solar analogs can be significantly different from the Sun, especially in the UV range, where the Sun can be variable by up to 20% (Colina et al. 1996). See Appendix B for more details.

In the following we use the expression

$$\frac{\varphi_{x,y\odot}}{\pi} = \frac{\int_0^{+\infty} \varphi_{\odot}(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}{\pi \cdot \int_0^{+\infty} M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda} (@1 \text{ AU}), \quad (10)$$

or

$$\frac{\varphi_{x,y\odot}}{\pi} = \frac{K_{x,y\odot}}{\pi \cdot \int_0^{+\infty} M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \frac{\lambda}{hc} \frac{\text{AP}}{\text{IG}} d\lambda},$$

to indicate the mean spectral irradiance of the Sun at a distance of 1 AU in the wavelength range defined by the passband  $(x, y)$ . Since the star observed by OSIRIS cannot be directly the Sun, but is a solar analog, the counts on the CCD refer to this second star. We can thus infer the “solar” counts from the analog counts ( $K_{x,y \text{ SA}}$ ) by using

$$K_{x,y\odot} = K_{x,y \text{ SA}} 10^{-0.4(V_{\odot}-V_{\text{SA}})},$$

so that the reflectance factor is

$$\frac{\varphi_{x,y\odot}}{\pi} = \frac{K_{x,y \text{ SA}} 10^{-0.4(V_{\odot}-V_{\text{SA}})}}{\pi \cdot \int_0^{+\infty} M^n(\lambda) \cdot F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \frac{\lambda}{hc} \frac{\text{AP}}{\text{IG}} d\lambda}, \quad (11)$$

**Table 6.** Selected UV calibrators from the BPGS catalog.

HIP	TDI	Name	Type	SpT	RA(ICRS)	Dec(ICRS)	Glon (deg)	Glat (deg)	Plx mas	mag <i>U</i>	mag <i>B</i>	mag <i>V</i>	m1565	m1965	m2365	m2740
88469		9 Sgr	Em*	O4V((f))	18 03 52.44466	-24 21 38.6269	6.0089	-01.2050	0.49	5.110	5.933	5.931				
97796	25 593	9 Sge	EB*	O7.5Ia...	19 52 21.76387	+18 40 18.7509	56.4824	-04.3314	-0.30	5.310	6.209	6.243	3.103	4.184	4.067	3.558
103371	27 499	HR 8023	*	O6V((f))	20 56 34.77848	+44 55 29.0048	85.6967	-00.2996	0.60	5.160	5.981	5.966	3.048	4.081	3.996	3.538
25980	4826	BD-01 935	V*	B1IV	05 32 41.35338	-01 35 30.5905	205.1376	-18.2024	2.09	4.240	5.150	5.340	1.187	2.827	2.255	1.800
103732	27 586	60 Cyg	Be*	B1Ve	21 01 10.92841	+46 09 20.7816	87.1548	-00.1028	2.14	4.230	5.217	5.417	1.427	2.924	2.327	1.967
88886	21 752	102 Her	*1*	B2IV	18 08 45.49142	+20 48 52.4079	47.4189	+18.4248	3.56	3.380	4.191	4.355	0.632	2.147	1.653	1.221
42799	12 946	$\eta$ Hya	bC*	B3V	08 43 13.47499	+03 23 55.1867	223.2510	+26.3204	5.56	3.380	4.100	4.300	0.595	2.120	1.555	1.132
86414	20 820	$\iota$ Her	bC*	B3IV	17 39 27.88609	+46 00 22.8001	72.3219	+31.2650	7.17	2.940	3.630	3.800	0.249	1.739	1.224	0.790
101909	27 077	HR 7899	*1*	B3V	20 39 04.96819	+15 50 17.5206	60.1618	-15.2527	2.42	5.110	5.820	5.980	2.479	3.917	3.477	3.041
84605	20 178	38 Oph	*1*	B9.5V	17 17 39.53024	-26 37 44.4788	358.5545	+06.4397	6.41		6.830	6.810	5.263	6.106	5.880	5.478
93104	23 639	HR 7174	Al*	B7IV	18 58 01.89799	+38 15 58.2810	68.5130	+15.2064	5.64		5.798	5.886	3.351	4.481	4.132	3.720
96275	24 977	9 Vul	V*	B8IIIIn	19 34 34.89569	+19 46 24.2462	55.3515	-00.1416	5.90	4.490	4.900	4.991	2.487	3.629	3.287	2.888
98441	25 869	HD 189689	Be*	B9	20 00 07.05730	+32 47 22.8336	69.4690	+01.4940	1.58		7.221	7.282	5.216	6.143	5.898	5.463
64238	16 666	$\theta$ Vir	*1*	A1IVs+...	13 09 56.99067	-05 32 20.4185	311.4176	+57.0274	10.33		4.381	4.381	3.611	4.007	3.883	3.514
100310	26 510	$\nu$ Cap	*1*	B9IV	20 20 39.81562	-12 45 32.6844	31.1700	-25.4782	12.88		4.716	4.754	3.204	4.058	3.786	3.420
97634	25 579	HR 7567	Al*	B1III+...	19 50 37.32728	+40 35 59.1378	75.2156	+07.1328	1.54	4.860	5.622	5.680	3.020	3.804	3.679	3.576
97774	25 640	HR 7591	V*	B2III	19 52 07.16659	+47 55 54.4611	81.8068	+10.5126	1.96	4.870	5.723	5.893	2.189	3.552	3.046	2.721
94385	24 126	20 Aql	V*	B3V	19 12 40.71241	-07 56 22.2610	28.2285	-08.3118	3.33		5.450	5.362	3.261	4.301	4.458	4.056
96491	25 114	HR 7467	*	B5II-III	19 36 56.64950	+38 23 01.7740	71.9693	+08.3674	2.77	5.720	6.374	6.498	3.452	4.699	4.326	3.907
93903	23 989	$\iota$ Lyr	Be*	B6IV	19 07 18.12878	+36 06 00.5654	67.2324	+12.6461	3.64	4.660	5.143	5.249	2.476	3.638	3.310	2.912
95067	24 509	HR 7346	*1*	B9V	19 20 33.05554	+35 11 09.4468	67.5696	+09.8229	3.47	5.810	6.196	6.287	3.724	4.896	4.488	4.089
83313	19 866	59 Her	*	A3IV	17 01 36.36149	+33 34 05.7613	56.1333	+36.4457	10.12		5.294	5.273	5.235	5.011	4.902	4.531
87192	21 062	HR 6642	*	A1V	17 48 47.87388	+19 15 18.0624	43.9685	+22.1582	6.04		6.154	6.132	4.742	5.501	5.336	4.937
98234	25 767	11 Sge	*	B9III	19 57 45.44583	+16 47 20.9792	55.5157	-06.3974	7.16	5.310	5.485	5.533	3.726	4.677	4.402	4.033
105811	28 096	69 Cyg	Pu*	B0Ib	21 25 47.02562	+36 40 02.5832	83.3947	-09.9588	0.36	4.890	5.860	5.940	2.623	3.608	3.270	2.947
99145	26 103	HR 7699	*	B5Ib	20 07 41.44099	+34 25 22.4644	71.6918	+01.0307	1.88	6.090	6.255	6.152	5.045	5.306	5.680	5.397
103312	27 484	HR 8020	Pu*	B8Iae	20 55 49.80412	+47 25 03.5627	87.5108	+01.4153	-0.22	5.640	6.094	5.727	6.232	5.726		6.443

**Notes.** Column description: HIP: HIPPARCOS number; TDI: TDI number; name: common name; type: object classification in Simbad; SpT: spectral type; RA, Dec: right ascension and declination; GLON, GLAT: Galactic coordinates; Plx: parallax; mag *U*, *B*, *V*: magnitudes in *U*, *B*, *V* filters in Johnson system. m1565, m1965, m2365, m2740: magnitudes from TDI survey at 1565, 1965, 2365, 2740 angstroms. Description of types. \*: star; \*1\*: star in double system; EB\*: eclipsing binary; Al\*: eclipsing binary of Algol type (detached); V\*: variable star; bC\*: variable Star of  $\beta$  Cep type; Pu\*: pulsating variable star; Be\*: Be star; Em\*: emission-line star.



where the use of the  $V$  magnitude, at each wavelength range, for the computation of the factor  $10^{-0.4(V_{\odot}-V_{SA})}$  is justified since we assume the same spectral shape for the Sun and the solar analog. Some authors refer to the value of the solar flux as already divided by  $\pi$ . To avoid confusion, we always explicitly write  $\varphi_{x,y\odot}/\pi$  as the solar intensity at 1 AU.

We can compute the reflectance directly from raw images (bias and flat-field corrected) by defining a composed absal reflectance factor, multiplying the absal factor defined in Eq. (3) and the new reflectance factor as expressed in Eq. (10).

Hence, the absal reflectance factor defined is the value needed to convert the count rate into the reflectance of a body located at 1 AU from the Sun:

$$A_{x,y} \cdot \frac{\varphi_{x,y}}{\pi} = \frac{K_{x,y \text{ calstar}} \cdot \text{pxsz}}{\pi} \times \frac{\int_0^{+\infty} \varphi_{\odot}(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}{\int_0^{+\infty} \varphi_{\text{calstar}}(\lambda) M^n(\lambda) F_{x,y}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda} \quad (@1 \text{ AU}), \quad (12)$$

where the “calstar” source (calibration star, usually Vega) used for the absolute flux calibration is generally different from the “SA” source (solar analog, usually 16 Cyg A+B, or 16 Cyg B) used to compute the solar value in Eq. (11). In Table 4 we report the absal reflectance factor values computed by means of Eq. (12).

To derive the value needed for a body at a heliocentric distance  $r$ , we simply have to divide this factor by  $r^2$  if  $r$  is expressed in AU.

#### 4.1. Counts from the solar analog

In the visible range the S/N of 16 Cyg B for the WAC is better than 100 with 180 s exposure time even with narrowband filters. For the NAC, the S/N is better than 200 with an exposure of a few seconds from about 480 nm to 1000 nm. In this spectral range we can also use this solar analog as a flux calibration star. Equation (12) can thus be further simplified by using “SA” as “calstar”, giving

$$A_{x,y} \cdot \frac{\varphi_{x,y\odot}}{\pi} = \frac{K_{x,y \text{ SA}} \times \text{pxsz} \times 10^{-0.4(V_{\odot}-V_{SA})}}{\pi} @1 \text{ AU}, \quad (13)$$

since we used  $\varphi_{\odot} = \varphi_{SA} 10^{-0.4(V_{\odot}-V_{SA})}$ .

This leads to smaller error bars because we simplified every instrumental passband; the passbands are the parameters carrying the largest uncertainties. Moreover, through this method we can exclude some problems that are due to the instrumental variation with time (age) and with some uncertainties in the optics efficiencies because they all affect the target and stellar images in the same way.

However, we have to remember that there may be some differences between the Sun and the solar analog, especially in the UV, where the calibration with both 16 Cyg A+B may lead to further approximations, and that this solar analog is quite faint in this wavelength range so that we have to work with a low S/N in the UV filters.

UV counts of 16 Cyg A+B have large error bars because the spectrum is very weak in these wavelengths. In this case, the only way to find a reflectance value is to return to Eq. (12), computing the count rate from a star like Vega (which is much brighter than the Sun, especially in the UV). The division by the weighted Vega spectrum and the multiplication by the weighted

Sun spectrum corrects the difference caused by the two different spectral types.

The two methods are expected to give the same values. If the absolute calibration is reliable and if the absal factors are the same regardless of the calibration star used, we can choose one of the two formulations to compute the best reflectance factor just according to the smallest error bars associated with the result.

Vega is a single star and has a much better  $S/N$  than 16 Cyg A or B in each filter. After checking that the calibration is consistent when computed using different calibrators, we found that the best star to use is Vega, both for the absal factor computation and for the solar values computation.

This formulation was verified to be valid. Good matches, within error bars, between OSIRIS spectrophotometry and literature reflectance spectra can be found for Phobos (Pajola et al. 2012, 2013) and Lutetia (Magrin et al. 2012).

#### 4.2. Twin combinations

When we have to calibrate NAC  $F84$  images from NAC  $F24$  baseline acquisition, we have to substitute Eq. (4) into Eq. (12), finding

$$A_{84} \cdot \frac{\varphi_{84\odot}}{\pi} = \frac{K_{24 \text{ calstar}} \cdot \text{pxsz}}{\pi} \times \frac{\int_0^{+\infty} \varphi_{\odot}(\lambda) \cdot M^n(\lambda) \cdot F_{84}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda}{\int_0^{+\infty} \varphi_{\text{calstar}}(\lambda) \cdot M^n(\lambda) \cdot F_{24}(\lambda) T_{\text{ARP}}(\lambda) Q(\lambda) \lambda d\lambda} @1 \text{ AU}.$$

This method was applied to complete the absolute reflectance calibration presented in Table 4 for filters without direct observation of the calibration stars.

### 5. Conclusions

We detected a good stability within  $\pm 1.5\%$  of the instrumental performances, without evidence of aging effects on the optics from 2007 to 2010. We found a good match between observed and expected data as deduced from on-ground laboratory measured optics and CCD efficiencies in the visible range. In the UV and IR wavelength range, however, we noted lower instrumental efficiencies than expected.

A reliable formulation for the flux calibration was proposed to be able to monitor this behavior from just after the exit from hibernation of the cameras and during the whole journey toward and around the comet, to study the possible dust contamination of the optics that would imply a wavelength-dependent variation of optical efficiencies. We also presented the formulation for the reflectance calibration. This formulation was verified to be valid, since there are good matches between OSIRIS spectrophotometry and literature reflectance spectra for Phobos (Pajola et al. 2012, 2013) and Lutetia (Magrin et al. 2012).

While the calibration compensates for instrumental efficiency variations in the visible and IR range, a deeper investigation of UV filters must be implemented. For this purpose, we finally suggested a list of other calibration sources. This will ensure a good calibration across the whole wavelength range during the entire last part of the mission.

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**Table B.1.** Properties of the Sun, 16 Cyg A, and 16 Cyg B according to Takeda (2005) and Takeda et al. (2005).

	Star	Sp. type	$T_{\text{eff}}$	$V$	$\pi$	$M_V$	B.C.	$\log(L/L_\odot)$	$L/L_\odot$	mass
	Sun	G2V	5778	-26.74		4.83				1
	HD 186408 16 Cyg A	G1.5V	5765.1	5.99	46.2	4.32	-0.13	+0.24	1.74	1.02
	HD 186427 16 Cyg B	G2.5V	5795.4	6.25	46.7	4.60	-0.12	+0.10	1.26	1.03

**Notes.** The HD number and the conventional stellar name for the three stars; the spectral type; the effective temperature derived from the spectroscopic approach reported in Takeda et al. (2005); the apparent  $V$  magnitude and the trigonometric parallax  $\pi$  in milliarcsecond, taken from the HIPPARCOS Catalog (ESA 1997); the absolute  $V$  magnitude,  $M_V$ , derived from the  $\pi$  and  $V$  values; the bolometric correction obtained from the results reported in Alonso et al. (1995); stellar luminosity  $\log(L/L_\odot)$  and  $L/L_\odot$  derived in Takeda et al. (2005) from  $M_V$  and B.C.; stellar mass (in units of  $M_\odot$ ) determined by  $L/L_\odot$  and  $T_{\text{eff}}$ .

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## Appendix A: Units

In the following we list and define the physical parameters and units we used.

- The irradiance is the power incident on a surface. It is expressed in the International System as  $\text{W m}^{-2}$ .
- The spectral irradiance is the power incident on a surface per unit wavelength, here expressed as  $\text{W m}^{-2} \text{nm}^{-1}$ .
- The radiance is the power emitted by a source per unit solid angle and unit area perpendicular to the emitting direction, expressed as  $\text{W m}^{-2} \text{sr}^{-1}$ .
- The spectral radiance is the power per unit solid angle per unit area per unit wavelength, here expressed as  $\text{W m}^{-2} \text{nm}^{-1} \text{sr}^{-1}$ .

The abscl factor presented in this paper is the factor needed to convert images from count rate into spectral radiance, so that the pixel-per-pixel signal associated with the target, observed in a particular passband, is directly comparable with a flux calibrated spectrum of the target in a wavelength interval around the central wavelength of the band.

## Appendix B: Description of the double star 16 Cyg A+B

The two stars 16 Cyg A and B are slightly different from the Sun. In particular, 16 Cyg A is larger than 16 Cyg B, and both are larger than the Sun. From Takeda (2005) and Takeda et al. (2005) we can find the  $\log L$  of both 16 Cyg A and B, in terms of  $L_\odot$ , transformed from the absolute visual magnitude  $M_V$  by means of the bolometric correction (B.C.) obtained by interpolating the results reported in Alonso et al. (1995). We summarize these results in Table B.1, where  $T_{\text{eff}}$  is computed for all stars from a spectroscopic method using the equivalent width values of FeI and FeII lines. Both 16 Cyg A and B are very close to the Sun.

If placed at the same distance as the Sun, the flux of 16 Cyg B is about 1.26 times the solar flux, and the flux of the pair 16 Cyg A+B is about three times the solar flux.

In summary: 16 Cyg A and B are slightly different from the Sun, but they seem to be similar enough to the Sun to allow us to assume the same spectral shape for the three stars, at least to the first order. Then, the simplest way to compute the scaling factor to convert the flux of 16 Cyg A+B into the solar flux is to consider  $V_{16 \text{ Cyg A}} = 5.99$ ,  $V_{16 \text{ Cyg B}} = 6.25$ , and  $V_\odot = -26.74$  at 1 AU. The magnitude  $V$  of the pair is then

$$V_{16 \text{ Cyg A+B}} = -2.5 \cdot \log_{10} 10^{-0.4 \cdot V_{16 \text{ Cyg A}}} + 10^{-0.4 \cdot V_{16 \text{ Cyg B}}}$$

and the difference between the magnitude  $V$  of the 16 Cyg pair and the Sun leads to the relation

$$\varphi_{16 \text{ Cyg A+B}} = \varphi_\odot \cdot 10^{-0.4 \cdot (V_{16 \text{ Cyg AB}} - V_\odot)}$$

because we assume this factor to be a constant along the spectrum, since we assume the same spectral shape for the three stars.

To estimate the count rate of the Sun from the count rate of 16 Cyg A+B, the relation is

$$K_{x,y \odot} = K_{x,y 16 \text{ Cyg A+B}} / 10^{-0.4 \cdot (V_{16 \text{ Cyg AB}} - V_\odot)}$$

at 1 AU, as we wrote in previous paragraphs.

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