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J. -F. Donati, E. Hébrard, G. A. J. Hussain, C. Moutou, L. Malo, et al.. Magnetic activity and hot Jupiters of young Suns: the weak-line T Tauri stars V819 Tau and V830 Tau. *Monthly Notices of the Royal Astronomical Society*, 2015, 453 (4), pp.3706–3719. 10.1093/mnras/stv1837 . hal-01439780

HAL Id: hal-01439780

<https://hal.science/hal-01439780v1>

Submitted on 4 Jan 2025

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Magnetic activity and hot Jupiters of young Suns: the weak-line T Tauri stars V819 Tau and V830 Tau

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Accepted 2015 August 6. Received 2015 August 5; in original form 2015 May 8

ABSTRACT

We report results of a spectropolarimetric and photometric monitoring of the weak-line T Tauri stars (wTTSs) V819 Tau and V830 Tau within the MaTYSSSE (Magnetic Topologies of Young Stars and the Survival of close-in giant Exoplanets) programme, involving the ESPaDOnS spectropolarimeter at the Canada–France–Hawaii Telescope. At ≈ 3 Myr, both stars dissipated their discs recently and are interesting objects for probing star and planet formation. Profile distortions and Zeeman signatures are detected in the unpolarized and circularly polarized lines, whose rotational modulation we modelled using tomographic imaging, yielding brightness and magnetic maps for both stars. We find that the large-scale magnetic fields of V819 Tau and V830 Tau are mostly poloidal and can be approximated at large radii by 350–400 G dipoles tilted at $\approx 30^\circ$ to the rotation axis. They are significantly weaker than the field of GQ Lup, an accreting classical T Tauri star (cTTS) with similar mass and age which can be used to compare the magnetic properties of wTTSs and cTTSs. The reconstructed brightness maps of both stars include cool spots and warm plages. Surface differential rotation is small, typically ≈ 4.4 times smaller than on the Sun, in agreement with previous results on wTTSs. Using our Doppler images to model the activity jitter and filter it out from the radial velocity (RV) curves, we obtain RV residuals with dispersions of 0.033 and 0.104 km s⁻¹ for V819 Tau and V830 Tau, respectively. RV residuals suggest that a hot Jupiter may be orbiting V830 Tau, though additional data are needed to confirm this preliminary result. We find no evidence for close-in giant planet around V819 Tau.

Key words: stars: magnetic fields – techniques: polarimetric – stars: formation – stars: imaging – stars: individual: V819 Tau – stars: individual: V830 Tau.

1 INTRODUCTION

In the last decades, our understanding of the key role that magnetic fields play during the first stages of the life of stars and their planets improved significantly, through both theoretical and observational advances (e.g. André, Basu & Inutsuka 2009). This is true in particular for low-mass pre-main-sequence (PMS) stars in the T Tauri star (TTS) phase, aged 0.5–10 Myr, that recently emerged from their dust cocoons and are still in a phase of gravitational contraction towards the main sequence (MS). Being either classi-

cal T-Tauri stars (cTTSs) when still surrounded by and accreting from massive, planet-forming accretion discs, or weak-line T-Tauri stars (wTTSs) once they have exhausted the gas in their inner discs, TTSs have been the subject of intense scrutiny at all wavelengths in recent decades given their interest for benchmarking the scenarios currently invoked to explain low-mass star and planet formation.

Magnetic fields of cTTSs are known to influence and sometimes even trigger several of the main physical processes playing a role at this stage, like accretion, outflows and angular momentum transport. As a result, they largely control the rotational evolution of low-mass PMS stars (e.g. Bouvier et al. 2007; Frank et al. 2014). More specifically, large-scale fields of cTTSs can evacuate the central regions of accretion discs, funnel the disc material on to the stars, and enforce

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corotation between cTTSs and their inner-disc Keplerian flows. This is causing cTTSs to rotate much more slowly than expected from the contraction and accretion of the high-angular-momentum disc material (e.g. Davies, Gregory & Greaves 2014). Fields of cTTSs and of their discs can also affect the formation and migration of planets (e.g. Baruteau et al. 2014). Last but not least, magnetic fields of cTTSs/wTTSs are known to trigger thermally driven winds through the heating provided by accretion shocks and/or Alfvén waves (e.g. Cranmer 2009; Cranmer & Saar 2011), leading to flares, coronal-mass ejections and angular momentum loss (e.g. Aarnio, Matt & Stassun 2012; Matt et al. 2012).

For these reasons, characterizing magnetic fields of cTTSs and wTTSs through observations is a crucial step to guide theoretical models towards more physical realism and reliable predictions. Although first detected more than 15 years ago (e.g. Johns-Krull, Valenti & Koresko 1999; Johns-Krull 2007), magnetic fields of TTSs are not yet fully characterized. In particular, field topologies of cTTSs have only recently been unveiled for a dozen stars (e.g. Donati et al. 2007; Hussain et al. 2009; Donati et al. 2010, 2013) through the MaPP (Magnetic Protostars and Planets) Large Observing Programme allocated on the 3.6-m Canada–France–Hawaii Telescope (CFHT) with the ESPaDOnS high-resolution spectropolarimeter (550 h of clear time over semester 2008b to 2012b). This first survey revealed, for instance, that large-scale fields of cTTSs remain relatively simple and mainly poloidal when the host star is still fully or largely convective, but become much more complex when the star turns mostly radiative (Gregory et al. 2012; Donati et al. 2013). This survey also demonstrated that these fields vary over a few years (e.g. Donati et al. 2011, 2012, 2013) and resemble those of mature stars with comparable internal structures (Morin et al. 2008b), hinting at a dynamo rather than fossil origin.

The situation is even worse for wTTSs, with only two of them magnetically imaged to date (namely V410 Tau and LkCa 4; Skelly et al. 2010; Donati et al. 2014). In both cases, the large-scale fields found are unexpected with respect to what we already know on cTTSs and MS dwarfs with similar internal structures. The most recent observations, focused on the wTTS LkCa 4, revealed the presence of a ring of intense toroidal field encircling the star at low latitudes, despite LkCa 4 being most likely fully convective and thus not supposed to exhibit this kind of magnetic feature (Donati et al. 2014). Clearly, a more systematic exploration of large-scale fields of wTTSs needs to be carried out in a similar manner to that of cTTSs. This is the goal of the new MaTYSSSE (Magnetic Topologies of Young Stars and the Survival of close-in giant Exoplanets) Large Programme, allocated at CFHT over semesters 2013a–2016b (510 h) with complementary observations with the NARVAL spectropolarimeter on the 2-m *Télescope Bernard Lyot* at Pic du Midi in France (420 h, allocated) and with the HARPS spectropolarimeter at the 3.6-m ESO Telescope at La Silla in Chile (70 h, allocated). MaTYSSSE also allows us to study magnetic winds of wTTSs and the corresponding spin-down rates (e.g. Vidotto et al. 2009, 2010, 2014); it also offers a novel option for assessing with improved sensitivity the potential presence of hot Jupiters (hJs) at an early stage in the life of low-mass stars and their planets, and thus for verifying whether core accretion coupled to migration is indeed the best option for forming close-in giant planets.

In practice, MaTYSSSE aims at carrying out a survey of about 30 wTTSs, along with a long-term monitoring of $\simeq 5$ cTTSs for which temporal variability of the large-scale magnetic field was already detected with MaPP. Rather than publishing the whole MaTYSSSE data set as a single study (which would imply both long delays to collect the data and carry out the analysis, and an

extremely long paper), our publication plan is to slice the work into smaller steps, by dedicating each forthcoming paper to only a few stars with similar properties. This ensures in particular that a significant fraction of observations can be published quickly and before data become public, and to advertise the MaTYSSSE results so that additional observing time (e.g. on HARPS-Pol) can be obtained in the meantime through complementary programmes. The statistical properties of the whole sample will be presented and discussed in a final wrap-up paper.

Following our initial MaTYSSSE paper (Donati et al. 2014), we present here a full analysis of the phase-resolved spectropolarimetric data sets collected on two new wTTSs, namely V819 Tau and V830 Tau. Both are well-studied single wTTSs (from direct imaging and spectroscopic monitoring; e.g. Kraus et al. 2011; Nguyen et al. 2012) and bona fide members of the Taurus L1495 dark cloud (Xiao et al. 2012), showing clear photometric variations of large amplitudes (Grankin et al. 2008; Xiao et al. 2012). V819 Tau is also known to exhibit a weak infrared excess indicating gas-poor remnants of a primordial disc or the early presence a warm debris disc (Furlan et al. 2009; Cieza et al. 2013).

We start our paper by documenting our observations (Section 2) and reviewing the spectral characteristics and evolutionary status of both stars (Section 3); we then describe the results we obtain by applying our tomographic modelling technique to these new data (Section 4), use them to filter out the activity jitter from RV curves and look for the potential presence of hJs around both stars (Section 5). Finally, we summarize our main results and discuss their implications for our understanding of low-mass star and planet formation (Section 6).

2 OBSERVATIONS

V819 Tau and V830 Tau were both observed in 2014 December and 2015 January using ESPaDOnS at the CFHT. ESPaDOnS collects stellar spectra spanning the entire optical domain (from 370 to 1000 nm) at a resolving power of 65 000 (i.e. a resolved velocity element of 4.6 km s^{-1}) over the full wavelength range (Donati 2003). A total of 15 circularly polarized (Stokes V) and unpolarized (Stokes I) spectra were collected for both stars over a time-span of 28 nights, corresponding to about five rotation cycles for V819 Tau and 10 cycles for V830 Tau. Time sampling was irregular in the first two thirds of the run (as a result of poor weather) but denser in the last third, yielding a reasonably good phase coverage for both stars despite a few remaining phase gaps of 0.15–0.20 cycle.

All polarization spectra consist of four individual subexposures taken in different polarimeter configurations to allow the removal of all spurious polarization signatures at first order. All raw frames are processed as described in the previous papers of the series (e.g. Donati et al. 2010, 2011, 2014), that the reader is referred to for more information. The peak signal-to-noise ratios (S/N; per 2.6 km s^{-1} velocity bin) achieved on the collected spectra range between 140 and 220 (median 190) for V819 Tau, and between 140 and 180 (median 170) for V830 Tau, depending mostly on weather/seeing conditions. All spectra are automatically corrected of spectral shifts resulting from instrumental effects (e.g. mechanical flexures, temperature or pressure variations) using telluric lines as a reference. Though not perfect, this procedure provides spectra with a relative RV precision of better than 0.030 km s^{-1} (e.g. Moutou et al. 2007; Donati et al. 2008). The full journal of observations is presented in Table 1 for both stars.

Rotational cycles of V819 Tau and V830 Tau (noted E_1 and E_2 in the following equation) are computed from Barycentric Julian

Table 1. Journal of ESPaDOnS observations of V819 Tau (first 15 lines) and V830 Tau (last 15 lines) collected from mid-2014 December to mid-2015 January. Each observation consists of a sequence of four subexposures, each lasting 1400 s and 700 s for V819 Tau and V830 Tau, respectively (except on 2015 Jan 15 for which subexposures on V819 Tau were shortened to 900 s). Columns 1–4 respectively list (i) the UT date of the observation, (ii) the corresponding UT time (at mid-exposure), (iii) the BJD, and (iv) the peak S/N (per 2.6 km s^{-1} velocity bin) of each observation. Column 5 lists the rms noise level (relative to the unpolarized continuum level I_c and per 1.8 km s^{-1} velocity bin) in the circular polarization profile produced by LSD, while Column 6 indicates the rotational cycle associated with each exposure (using the ephemerides given by equation 1).

Date (2014)	UT (hh:mm:ss)	BJD (2457 000+)	S/N	σ_{LSD} (0.01 per cent)	Cycle
Dec 19	12:07:54	11.01081	140	3.9	0.002
Dec 20	07:55:28	11.83548	210	2.5	0.151
Dec 21	07:17:08	12.80882	210	2.4	0.327
Dec 22	08:14:51	13.84886	200	2.6	0.515
Dec 28	10:50:17	19.95655	190	2.8	1.619
Dec 29	07:37:39	20.82273	180	2.9	1.776
Dec 30	06:10:02	21.76183	200	2.6	1.946
Jan 07	07:19:42	29.80976	190	2.7	3.401
Jan 08	06:20:11	30.76836	190	2.7	3.574
Jan 09	06:15:12	31.76484	190	2.9	3.754
Jan 10	06:15:60	32.76533	170	3.2	3.935
Jan 11	07:25:26	33.81348	220	2.3	4.125
Jan 12	06:28:22	34.77379	200	2.4	4.298
Jan 14	06:17:48	36.76631	190	2.8	4.658
Jan 15	06:12:46	37.76274	150	3.5	4.839
Dec 20	09:13:43	11.88992	170	2.8	0.033
Dec 21	08:33:48	12.86217	170	2.8	0.387
Dec 22	09:29:48	13.90104	180	2.7	0.766
Dec 28	12:19:59	20.01897	140	3.6	2.999
Dec 29	08:53:59	20.87588	160	3.0	3.311
Dec 30	07:26:53	21.81535	160	2.9	3.654
Jan 07	08:35:55	29.86285	170	2.8	6.590
Jan 08	07:36:48	30.82174	180	2.7	6.940
Jan 09	07:31:38	31.81810	170	2.9	7.303
Jan 10	07:32:25	32.81858	150	3.3	7.668
Jan 11	08:42:02	33.86686	180	2.6	8.051
Jan 12	05:12:50	34.72151	170	2.9	8.362
Jan 13	05:03:32	35.71500	160	3.0	8.725
Jan 14	05:02:24	36.71414	170	2.8	9.089
Jan 15	07:12:21	37.80431	170	2.9	9.487

Dates (BJDs) according to the ephemerides:

$$\begin{aligned} \text{BJD(d)} &= 2457011.0 + 5.53113E_1 \text{ (for V819 Tau)} \\ \text{BJD(d)} &= 2457011.8 + 2.74101E_2 \text{ (for V830 Tau)} \end{aligned} \quad (1)$$

in which the photometrically determined rotation periods P_{rot} (equal to 5.53113 and 2.74101 d, respectively; Grankin 2013) are taken from the literature and the initial Julian dates (2457011.0 and 2457011.8 d) are chosen arbitrarily.

Least-Squares Deconvolution (LSD; Donati et al. 1997) was applied to all spectra. The line list we employed for LSD is computed from an ATLAS9 LTE model atmosphere (Kurucz 1993) featuring an effective temperature of $T_{\text{eff}} = 4250 \text{ K}$ and a logarithmic gravity (in cgs unit) of $\log g = 4.0$, appropriate for both V819 Tau and V830 Tau (see Section 3). As usual, only moderate to strong atomic spectral lines are included in this list (see e.g. Donati et al. 2010, for more details). Altogether, about 7800 spectral features (with about 40 per cent from Fe I) are used in this process. Expressed in units of the unpolarized continuum level I_c , the average noise levels of the

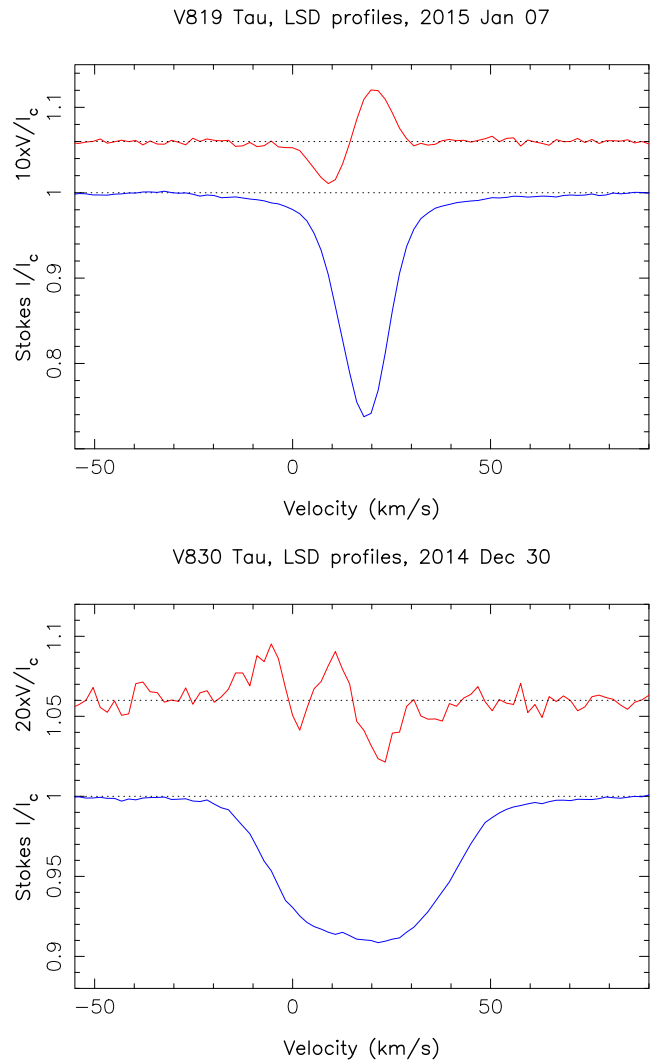


Figure 1. LSD circularly polarized (Stokes V , top/red curve) and unpolarized (Stokes I , bottom/blue curve) profiles of V819 Tau (top panel) and V830 Tau (bottom panel) collected on 2015 Jan. 07 (cycle 3.401) and 2014 Dec. 30 (cycle 3.654). Clear Zeeman signatures are detected in the LSD Stokes V profile of both stars (with a complex shape in the case of V830 Tau), in conjunction with the unpolarized line profiles. The mean polarization profiles are expanded (by a factor of 10 and 20 for V819 Tau and V830 Tau, respectively) and shifted upwards by 1.06 for display purposes.

resulting Stokes V LSD signatures range from 2.3 to 3.9×10^{-4} per 1.8 km s^{-1} velocity bin – with a median value of 2.8×10^{-4} for both stars.

Zeeman signatures are detected at all times in Stokes V LSD profiles (see Fig. 1 for an example), featuring amplitudes of 0.3–1 per cent, i.e. indicative of significant large-scale fields. Asymmetries and/or distortions are also visible in Stokes I LSD profiles, suggesting the presence of brightness inhomogeneities at the surfaces of V819 Tau and V830 Tau at the time of our observations.

Contemporaneous BVR_J photometric observations were also collected from the Crimean Astrophysical Observatory (CrAO) 1.25-m telescope between 2014 Aug and Dec for both stars, indicating that V819 Tau and V830 Tau were exhibiting brightness modulations with full amplitudes of 0.45 and 0.10 mag in V (see Table 2) and periods of 5.525 ± 0.015 and 2.741 ± 0.005 d, respectively (compatible within error bars with the average periods of Grankin 2013,

Table 2. Journal of contemporaneous CrAO multi-colour photometric observations of V819 Tau (first 22 lines) and V830 Tau (last 19 lines) collected in late 2014, respectively, listing the UT date and Heliocentric Julian Date (HJD) of the observation, the measured V magnitude, $B - V$ and $V - R_J$ Johnson photometric colours, and the corresponding rotational phase (using again the ephemerides given by equation 1).

Date (2013)	HJD (2456 000+)	V (mag)	$B - V$	$V - R_J$	Phase
Aug 26	896.5665	12.824	1.512	1.461	0.311
Aug 27	897.5395	12.954	1.488	1.478	0.487
Aug 29	899.5421	13.124	1.577	1.520	0.849
Aug 30	900.5427	12.911	1.486	1.522	0.030
Aug 31	901.5438	12.788	1.522	1.487	0.211
Sept 01	902.5487	12.922	1.524	1.509	0.393
Sept 02	903.5425	13.154	1.549	1.614	0.572
Sept 04	905.5024	12.965	1.558	1.518	0.927
Sept 05	906.5354	12.801	1.549	1.470	0.113
Sept 20	921.5742	13.142	1.499	1.544	0.832
Sept 20	921.5742	13.204	1.519	1.522	0.832
Oct 05	936.6035	13.099	1.618	1.447	0.549
Oct 15	946.5991	12.898	1.496	1.541	0.357
Oct 26	957.4906	12.838	1.514	1.485	0.326
Oct 28	959.5578	13.259	1.572	1.578	0.700
Nov 02	964.4855	13.163	1.658	1.539	0.590
Nov 05	967.6033	12.784	1.656	1.437	0.154
Nov 05	967.6106	12.795	1.518	1.513	0.155
Nov 13	975.4293	13.159	1.611	1.574	0.569
Nov 14	976.4257	13.230	1.558	1.592	0.749
Dec 13	1005.4237	12.956	1.615	1.516	0.992
Dec 14	1006.5232	12.804	1.586	1.506	0.191
Aug 22	892.5484	12.321	1.351	1.319	0.494
Aug 27	897.5627	12.297	1.403	1.297	0.323
Aug 29	899.5574	12.209	1.370	1.270	0.051
Aug 30	900.5540	12.288	1.358	1.311	0.414
Aug 31	901.5514	12.238	1.368	1.296	0.778
Sept 01	902.5578	12.262	1.378	1.288	0.145
Sept 02	903.5502	12.321	1.375	1.351	0.507
Sept 04	905.5599	12.304	1.389	1.286	0.241
Sept 05	906.5443	12.289	1.387	1.298	0.600
Oct 05	936.5812	12.278	1.398	1.999	0.558
Oct 15	946.5610	12.297	1.405	1.303	0.199
Oct 28	959.5923	12.220	1.365	1.298	0.953
Nov 02	964.5193	12.274	1.387	1.275	0.751
Nov 05	967.6268	12.210	1.387	1.275	0.884
Nov 05	967.6307	12.224	1.345	1.299	0.886
Nov 13	975.4453	12.280	1.379	1.294	0.737
Nov 14	976.4437	12.277	1.351	1.321	0.101
Dec 13	1005.4898	12.319	1.367	1.326	0.698
Dec 14	1006.4869	12.268	1.403	1.309	0.062

used to phase our spectroscopic data; see equation 1). These photometric variations (unusually large for V819 Tau and small for V830 Tau; see fig. 12 a and table 3 of Grankin et al. 2008) are commonly attributed to the presence of brightness features at the surface of both stars.

3 EVOLUTIONARY STATUS OF V819 TAU AND V830 TAU

Applying to our highest S/N spectra the automatic spectral classification tool especially developed in the context of MaPP and MaTYSSSE, inspired from that of Valenti & Fischer (2005) and discussed in a previous paper (Donati et al. 2012), we find that the photospheric temperatures and logarithmic gravities of V819

Tau and V830 Tau are, respectively, equal, for both stars, to $T_{\text{eff}} = 4250 \pm 50$ K and $\log g = 3.9 \pm 0.2$ (with g in cgs units). Repeating the operation on different spectra of our data set yields consistent parameters, with differences between results within the quoted error bars. The temperature we obtain is higher than that often quoted in the literature for both stars (in the range 3900–4100 K; see e.g. Sestito, Palla & Randich 2008; Furlan et al. 2009; Grankin 2013; Herczeg & Hillenbrand 2014), and derived in most cases from either photometry or low-resolution spectroscopy. In the particular case of V830 Tau, our estimate is in agreement with the only other spectroscopic study based on high-resolution data that we know of in the refereed literature (Schiavon, Batalha & Barbuy 1995). Our result also confirms that V819 Tau and V830 Tau are warmer than our previous target LkCa 4 (whose temperature we estimated to be $T_{\text{eff}} = 4100 \pm 50$ K; Donati et al. 2014). From the difference between the $B - V$ index expected at this temperature (1.15 ± 0.02 , as estimated through an interpolation within table 6 of Pecaut & Mamajek 2013) and the averaged value measured for both V819 Tau and V830 Tau (equal to 1.55 ± 0.05 and 1.35 ± 0.05 , respectively; see Kenyon & Hartmann 1995; Grankin et al. 2008, see also Table 2), we derive that the amounts of visual extinction A_V that our two targets suffer are equal to 1.24 ± 0.15 for V819 Tau and 0.62 ± 0.15 for V830 Tau (in reasonable agreement with other independent results; e.g. Herczeg & Hillenbrand 2014).

Long-term photometric monitoring of V819 Tau and V830 Tau indicates that their maximum brightnesses correspond to V magnitudes of 12.8 and 11.9, respectively (Grankin et al. 2008); assuming a spot coverage of $\simeq 25$ per cent of the visible stellar hemisphere at maximum brightness, typical for active stars (and caused by, e.g. the presence of high-latitude cool spots or by a network of small spots evenly spread in longitude), we derive unspotted V magnitudes of 12.5 ± 0.2 and 11.6 ± 0.2 for V819 Tau and V830 Tau, respectively. Using the visual bolometric correction expected for the adequate photospheric temperature (equal to -0.86 ± 0.05 ; Pecaut & Mamajek 2013) and the accurate distance estimate of the Taurus L1495 dark cloud (131 ± 3 pc, corresponding to a distance modulus equal to 5.59 ± 0.05 ; Torres et al. 2012), we finally obtain bolometric magnitudes of 4.81 ± 0.25 and 4.53 ± 0.25 , or equivalently logarithmic luminosities relative to the Sun of -0.03 ± 0.10 and 0.08 ± 0.10 for V819 Tau and V830 Tau, respectively. Coupling with the photospheric temperature obtained previously, we find respective radii of 1.8 ± 0.2 and $2.0 \pm 0.2 R_{\odot}$ for our two target stars.

The rotation periods of both V819 Tau and V830 Tau are well determined from long-term multi-colour photometric monitoring, and equal to 5.53113 d and 2.74101 d (Grankin 2013), respectively; estimates from different studies (e.g. Xiao et al. 2012) agree to a precision of better than 1 per cent, indicative of robust and reliable measurements. Coupling these rotation periods along with our measurements of the line-of-sight-projected equatorial rotation velocities $v \sin i$ of V819 Tau and V830 Tau (respectively equal to 9.5 ± 0.5 and 30.5 ± 0.5 km s $^{-1}$; see Section 4), we can infer that the $R_* \sin i$ values of both stars are $1.04 \pm 0.05 R_{\odot}$ and $1.65 \pm 0.03 R_{\odot}$, where R_* and i note the radius of the stars and the inclination of their rotation axis to the line of sight. Comparing with the radii derived from the luminosities and photospheric temperature, we derive that i is equal to 35° and 55° (to an accuracy of $\simeq 10^\circ$) for V819 Tau and V830 Tau, respectively.

Using the evolutionary models of Siess et al. (2000, assuming solar metallicity and including convective overshooting), we obtain that V819 Tau and V830 Tau are both $1.00 \pm 0.05 M_{\odot}$ stars, and

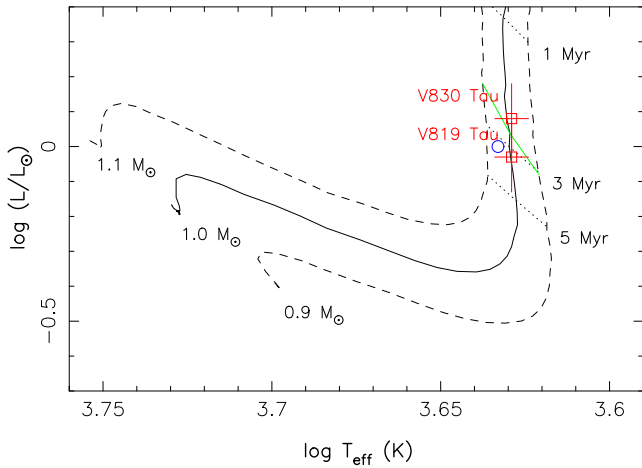


Figure 2. Observed (open squares and error bars) location of V819 Tau and V830 Tau in the HR diagram. The PMS evolutionary tracks (full and dashed lines) and corresponding isochrones (dotted lines, all from Siess, Dufour & Forestini 2000) assume solar metallicity and include convective overshooting. The blue circle indicates the location of the cTTS GQ Lup according to Donati et al. (2012). The green line depicts where models predict PMS stars start developing their radiative core as they contract towards the MS.

that their respective ages are $\simeq 3.2$ and $\simeq 2.2$ Myr (with relative error bars of $\simeq 50$ per cent; see Fig. 2), suggesting that V819 Tau may be a slightly evolved version (by about $\simeq 1$ Myr) of V830 Tau.¹ This supports that disc dissipation for PMS stars in Taurus may occur on time-scales as short as a few Myr (Williams & Cieza 2011; Ingleby et al. 2012), even though 80–90 per cent of single stars in this star formation region still host discs at similar ages (Kraus et al. 2012). This makes V819 Tau and V830 Tau atypical in this respect, and thus of obvious interest for MaTYSSSE. We also note that our two targets are both located close to the theoretical threshold at which $1 M_{\odot}$ stars cease to be fully convective (green line in Fig. 2). This may suggest that V830 Tau is still fully convective whereas V819 Tau has already developed a small radiative core; our error bars on the location of both stars in the HR diagram are, however, still too large to reach a firm conclusion (see Fig. 2), to the point that even the opposite conclusion could actually be true. Interestingly, V819 Tau and V830 Tau are also located close to the cTTS GQ Lup in the HR diagram (Donati et al. 2012, see circle in Fig. 2), implying that their internal structures should be similar, if that of GQ Lup is not significantly impacted by accretion (only moderate for this cTTS; e.g. Donati et al. 2012). We summarize the main parameters of both stars in Table 3.

Finally, we report that core emission is clearly detected in the Ca II infrared triplet (IRT) lines of both V819 Tau and V830 Tau, with an average equivalent width of the emission core equal to $\simeq 15 \text{ km s}^{-1}$, corresponding to the amount expected from chromospheric emission for such PMS stars. Moreover, the He I D_3 line is barely visible (average equivalent width of $\simeq 5 \text{ km s}^{-1}$) for both V819 Tau and V830 Tau, further demonstrating that accretion is no longer taking place on to their surfaces, in agreement with previous studies.

¹ Using the most recent evolutionary models of Baraffe et al. (2015), we obtain slightly smaller masses (of $0.90 M_{\odot}$) for both stars, along with smaller ages (of $\simeq 2.2$ and $\simeq 1.4$ Myr, respectively). We nevertheless kept estimates derived from Siess et al. (2000) as our reference values for internal consistency with previous MaPP and MaTYSSSE results.

Table 3. Main parameters of V819 Tau and V830 Tau as derived from our study except when indicated (with G13 and T12 standing for Grankin 2013 and Torres et al. 2012 respectively), with v_{rad} noting the RV that the star would have if unspotted (as inferred from the modelling of Section 4).

	V819 Tau	V830 Tau	Reference
$M_{\star} (M_{\odot})$	1.00 ± 0.05	1.00 ± 0.05	
$R_{\star} (R_{\odot})$	1.8 ± 0.2	2.0 ± 0.2	
age (Myr)	$\simeq 3.2$	$\simeq 2.2$	
$T_{\text{eff}} (K)$	4250 ± 50	4250 ± 50	
$\log(L_{\star}/L_{\odot})$	-0.03 ± 0.10	0.08 ± 0.10	
$P_{\text{rot}} (d)$	5.53113	2.74101	G13
$v \sin i (km s^{-1})$	9.5 ± 0.5	30.5 ± 0.5	
$v_{\text{rad}} (km s^{-1})$	16.6 ± 0.1	17.5 ± 0.1	
$i (^{\circ})$	35 ± 10	55 ± 10	
distance (pc)	131 ± 3	131 ± 3	T12

4 TOMOGRAPHIC MODELLING

Now that our two stars are well characterized regarding their atmospheric properties and evolutionary status, we are ready to apply our dedicated stellar-surface tomographic-imaging package to the spectropolarimetric data set described in Section 2. This tool is based on the principles of maximum-entropy image reconstruction and on the assumption that the observed variability is mainly caused by rotational modulation (with an added option for differential rotation); this framework is well adapted to the case of wTTSs for which rotational modulation largely dominates intrinsic variability, in both photometry (e.g. Grankin et al. 2008) and photospheric lines (e.g. Skelly et al. 2010; Donati et al. 2014). Since initially released (Brown et al. 1991; Donati & Brown 1997), the code underwent several upgrades (e.g. Donati 2001; Donati et al. 2006b), the most recent ones being its re-profiling to the specific needs of MaPP and MaTYSSSE observations (Donati et al. 2010, 2014). More specifically, the imaging code is set up to invert (both automatically and simultaneously) time series of Stokes I and V LSD profiles into brightness and magnetic maps of the stellar surface; moreover, brightness imaging is now allowed to reconstruct both cool spots and warm plages, known to contribute to the activity of very active stars (Donati et al. 2014). The reader is referred to the papers mentioned above for more general details on the imaging method.²

To compute the disc-integrated average photospheric LSD profiles, we start by synthesizing the local Stokes I and V profiles using Unno-Rachkovsky’s analytical solution to the polarized radiative transfer equations in a Milne-Eddington model atmosphere, taking into account the local brightness and magnetic field; we then integrate these local profiles over the visible hemisphere to retrieve the synthetic profiles to be compared with our observations. This computation method provides in particular a reliable description of how line profiles are distorted in the presence of magnetic fields (including magneto-optical effects; e.g. Landi degl’Innocenti & Landolfi 2004). The main parameters of the local profile are similar to those used in our previous studies, the wavelength, Doppler width, equivalent width and Landé factor being, respectively, set to 670 nm, 1.8 km s^{-1} , 3.9 km s^{-1} and 1.2. As part of the imaging process, we obtain accurate estimates for several parameters of both stars; we find in particular that the v_{rad} s (the RVs the stars would have if unspotted) and $v \sin i$ s are, respectively, equal to 16.6 ± 0.1 and

² As explained in Donati et al. (2014), full Stokes spectropolarimetry is not expected to bring much improvement in imaging quality over Stokes I and V spectropolarimetry, especially for stars as faint as our MaTYSSSE targets.

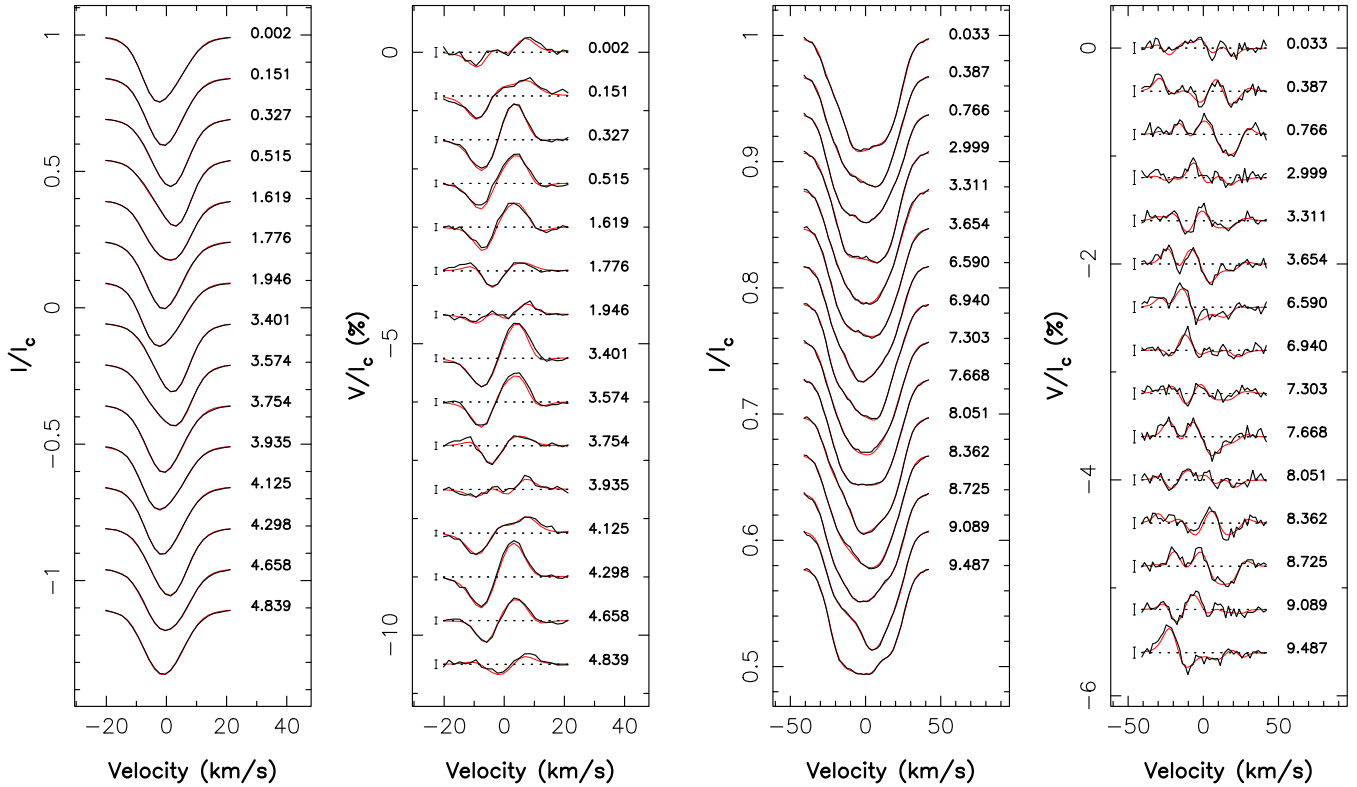


Figure 3. Maximum-entropy fit (thin red line) to the observed (thick black line) Stokes I (first and third panels) and Stokes V (second and fourth panels) LSD photospheric profiles of V819 Tau (first two panels) and V830 Tau (last two panels) in 2014 Dec and 2015 Jan. Rotational cycles and 3σ error bars (for Stokes V profiles) are also shown next to each profile. This figure is best viewed in colour.

$9.5 \pm 0.5 \text{ km s}^{-1}$ for V819 Tau, and 17.5 ± 0.1 and $30.5 \pm 0.5 \text{ km s}^{-1}$ for V830 Tau (as listed in Table 3), in good agreement with previously published estimates (e.g. Nguyen et al. 2012). We will revisit the RV curves of both stars in more details in Section 5.

4.1 Brightness and magnetic imaging

We show in Fig. 3 our sets of Stokes I and V LSD profiles of V819 Tau and V830 Tau along with our fits to the data. The fits we obtain correspond to a reduced chi-square χ_r^2 equal to 1 (i.e. a χ^2 equal to the number of fitted data points, respectively equal to 360 and 705 for V819 Tau and V830 Tau), starting from initial values of about 65 and 10 (corresponding to null magnetic fields and unspotted brightness maps) for V819 Tau and V830 Tau, respectively. This further stresses the quality of our data set and the high performance of our imaging code at modelling the observed modulation of LSD profiles (also obvious from Fig. 3). The very good quality of the achieved fit clearly demonstrates that the assumption underlying our modelling, i.e. that the observed line profile variability is caused by the presence of brightness and magnetic surface structures rotating in and out of the observer’s view, is verified. As for LkCa 4 (Donati et al. 2014), the reasonably dense phase coverage allows us to track back most of the main reconstructed features into genuine profile distortions, and thus to safely claim that our maps include no major imaging artefact or bias.

The reconstructed brightness maps of V819 Tau and V830 Tau (see Fig. 4) both include cool spots and warm plages, but nevertheless feature a number of significant differences, regarding the contrast of these features in particular, much larger for V819 Tau than for V830 Tau. This difference is readily visible from our pho-

tometric data (collected at CrAO; see Table 2 and Fig. 5) whose full amplitude is about 4.5 times larger ($\simeq 0.45 \text{ mag}$) for V819 Tau than for V830 Tau ($\simeq 0.1 \text{ mag}$). We emphasize that our reconstructed brightness images predict light curves that are in very good agreement with the observed ones (see Fig. 5), even though these images were produced solely from our sets of LSD profiles; this demonstrates that surface imaging from well-sampled series of high-quality spectroscopic data is capable of predicting photometric curves, at a typical relative precision of about 28 and 14 mmag for V819 Tau and V830 Tau, respectively.

The brightness map of V819 Tau (see Fig. 4, left panel) is reminiscent of that of LkCa 4 with a dark spot near the pole – even though this spot is not quite as dark and close to the pole as that of LkCa 4 (Donati et al. 2014); it also comes with a warm plage next to it though this plage is again less bright than its equivalent on LkCa 4. We suspect that part (but not all) of this difference reflects the lower $v \sin i$ of V819 Tau (with respect to LkCa 4) and thus the poorer spatial resolution achievable through Doppler imaging at the surface of this star. As for LkCa 4, bright plages are again required to reproduce all details of the profile variability observed for V819 Tau, and in particular the large RV variations (of full amplitude $\simeq 2.4 \text{ km s}^{-1}$; see Section 5) that V819 Tau exhibits. The reconstructed brightness features are found to cover $\simeq 15$ per cent of the overall stellar surface, with the main cool spot covering $\simeq 9$ per cent and the plage $\simeq 6$ per cent. Note that these estimates of the spot/plage coverage should be considered as lower limits only (hence the larger spot coverage assumed in Section 3 to derive the location of both stars in the HR diagram); Doppler imaging is indeed powerful at recovering large-scale features causing detectable profile distortions, but mostly insensitive to small-scale structures more or less

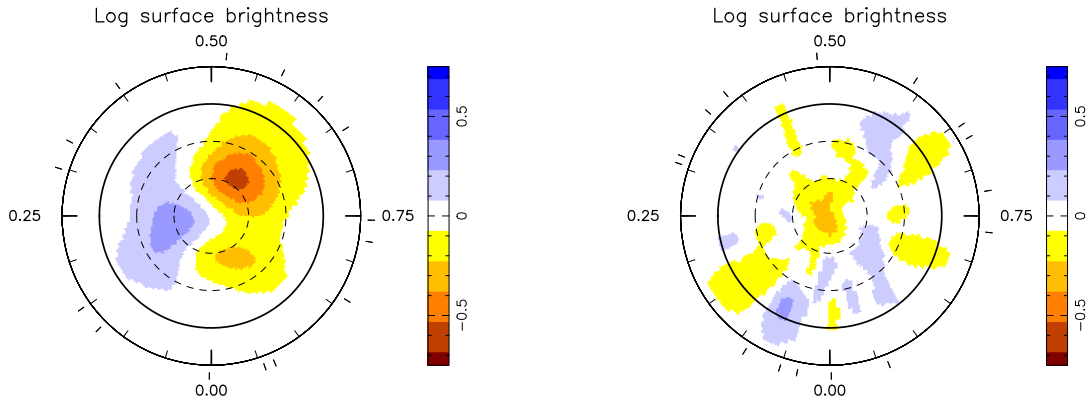


Figure 4. Maps of the logarithmic brightness (relative to the quiet photosphere), at the surfaces of V819 Tau (left) and V830 Tau (right) in 2014 Dec–2015 Jan. The stars are shown in flattened polar projection down to latitudes of -30° , with the equator depicted as a bold circle and parallels as dashed circles. Radial ticks around each plot indicate phases of observations. This figure is best viewed in colour.

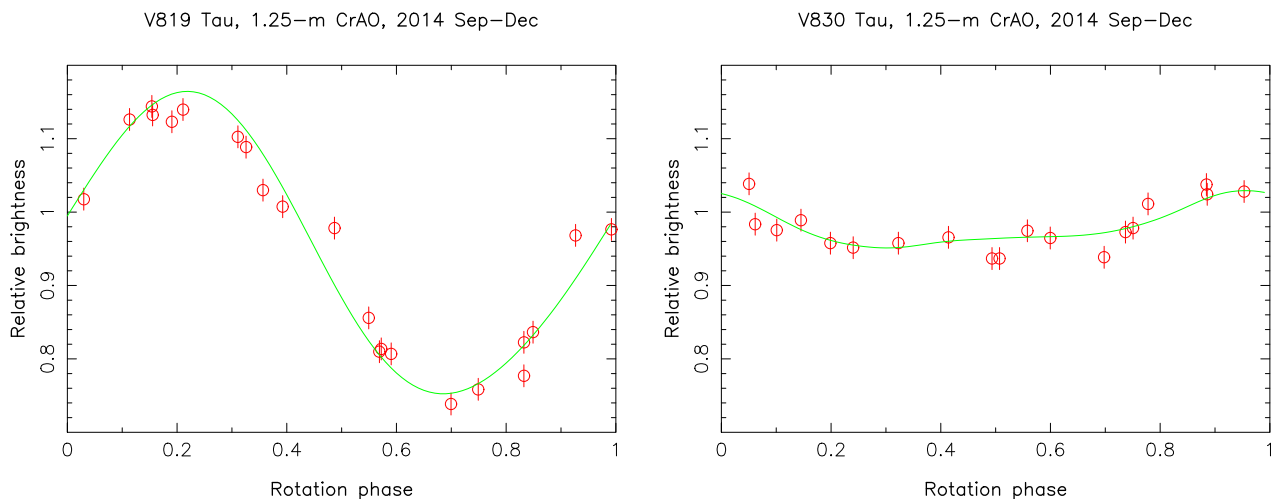


Figure 5. Brightness variations of V819 Tau (left) and V830 Tau (right) predicted from the tomographic modelling of Fig. 4 of our spectropolarimetric data set (green line), compared with contemporaneous photometric observations in the V band (red open circles and error bars of 15 mmag) at the 1.25-m CrAO telescope (see Table 2).

evenly distributed over the stellar surface and generating no profile distortions.

The brightness map of V830 Tau also features a cool spot close to the pole and at least one warm plage near the equator (see Fig. 4, right panel); although the overall surface covered with spots and plages on V830 Tau, equal to $\simeq 12$ per cent ($\simeq 6$ per cent for spots and $\simeq 6$ per cent for plages), is comparable to that of V819 Tau, they feature a lower contrast in average and are spread over a more even range of longitudes, yielding a much flatter photometric response (see Fig. 5, right panel) despite significant distortions of the LSD Stokes I profiles (see Fig. 3, third panel) and clear RV fluctuations (of full amplitude $\simeq 1.6$ km s $^{-1}$; see Section 5). As already evidenced in the case of V410 Tau (the first wTTS to be magnetically imaged, although with a slightly older version of our imaging code; Skelly et al. 2010), we confirm that a flat photometric curve by no means implies the absence of spots (or the reduction in spot coverage) at the surface of the star, but rather a more even distribution of features (both cool and warm). Moreover, the similar $v \sin i$ s of V830 Tau and LkCa 4 further argue (along with the different amplitudes of the photometric light curves) that the reconstructed brightness maps of these two stars trace genuine differences in the parent spot distributions at the surface of these two stars.

Finally, we emphasize that the good agreement obtained for both stars between the predicted and observed light curves despite the large time-span (of $\simeq 4$ months) over which our photometric data were collected (as opposed to $\simeq 1$ month only for our spectropolarimetric data; see Tables 1 and 2) brings further evidence that surface features of wTTSs are stable and long-lived on a time-scale of months, supporting previous conclusions from long-term photometric monitoring (e.g. Grankin et al. 2008).

The reconstructed magnetic fields of V819 Tau and V830 Tau (see Fig. 6) are found to be similar in terms of their intensities and topological properties, and are much weaker than the one found at the surface of LkCa 4 (featuring kG poloidal and toroidal field components; Donati et al. 2014). In both cases, the fields are mainly poloidal, at a level of 80 per cent and 90 per cent for V819 Tau and V830 Tau, respectively, with average unsigned fluxes of 370 and 300 G. Most of the reconstructed poloidal field energy (75 and 60 per cent) concentrates in spherical harmonics (SH) dipolar modes (i.e. with $\ell = 1$, ℓ denoting the degree of the modes), whereas 90 per cent of it gathers in the aligned dipole mode ($\ell = 1$ and $m = 0$, with m denoting the mode order) for both stars. At first order and at some distance from the stars, the poloidal components of V819 Tau and V830 Tau can be approximated with dipoles of

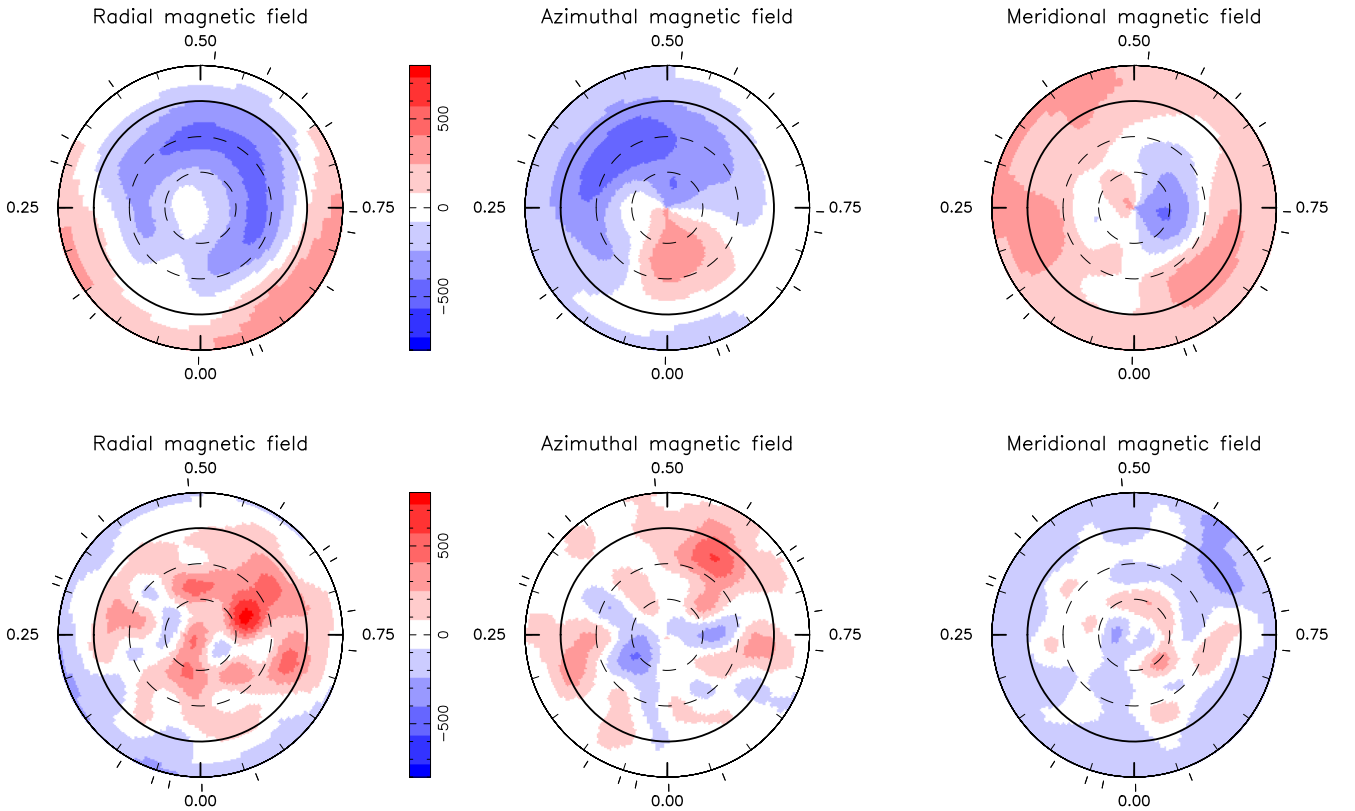


Figure 6. Maps of the radial (left), azimuthal (middle) and meridional (right) components of the magnetic field B at the surfaces of V819 Tau (top) and V830 Tau (bottom) in 2014 Dec–2015 Jan. Magnetic fluxes in the colour lookup table are expressed in G. The stars are shown in flattened polar projection as in Fig. 4. This figure is best viewed in colour.

respective strengths 400 and 350 G, tilted at angles of $\approx 30^\circ$ to the line of sight towards phases 0.50 and 0.65.³ A weaker octupolar component in the range 150–300 G (depending on whether or not we favour odd modes in the imaging process, as for cTTSs; Donati et al. 2011, 2012, 2013) is also present on V819 Tau, more or less aligned with the rotation axis and antiparallel with the dipole component; on V830 Tau, the octupolar component is weaker and non-axisymmetric, with a polar strength no larger than 150 G. The reconstructed large-scale fields of V819 Tau and V830 Tau also include weaker (though significant) toroidal components with average unsigned fluxes equal to ≈ 170 G and ≈ 100 G, respectively, and whose topologies are more complex and less axisymmetric than that of LkCa 4 (Donati et al. 2014).

SH expansions describing the reconstructed field presented in Fig. 6 are limited to terms with $\ell \leq 8$ and 10 for V819 Tau and V830 Tau, respectively; only marginal changes to the solution are observed when increasing the maximum ℓ value beyond these thresholds, demonstrating that most of the Zeeman signal detected in Stokes V LSD profiles of both stars is captured in the images we recovered. As an example, we present in Fig. 7 the extrapolated large-scale field topologies of V819 Tau and V830 Tau using the potential field approximation (e.g. Jardine, Collier Cameron & Donati 2002) and derived solely from the reconstructed

radial field components; as states of lowest possible magnetic energy, these potential topologies are shown to provide a reliable description of the magnetic field well within the Alfvén radius (Jardine et al. 2013). From these plots, the dominantly dipolar topologies of the large-scale fields of V819 Tau and V830 Tau are quite obvious.

Finally, we note that the reconstructed brightness (see Fig. 4) and magnetic maps (see Fig. 6) of V819 Tau and V830 Tau only show a weak level of spatial correlation, as opposed to what we reported for LkCa 4 where the cool polar spot overlapped almost exactly with the region of strongest radial field (see fig. 4 of Donati et al. 2014). The closest match of this kind that we can find in our new data is the cool spot close to the pole of V819 Tau, which roughly coincides with the visible pole of the dominant dipolar component of the large-scale field, but not with a local maximum of the surface field (as observed on LkCa 4). This is in line with what is observed on other cool active stars with large-scale field strengths similar to those of V819 Tau and V830 Tau (e.g. Donati et al. 2003a); observations so far suggest that such tight spatial correlation between brightness and magnetic maps is only observed for young stars in which the large-scale poloidal field exceeds 1 kG (like LkCa 4, but also e.g. V410 Tau, GQ Lup, V2129 Oph and DN Tau, Skelly et al. 2010; Donati et al. 2011, 2012, 2014).

4.2 Surface differential rotation

Spread out over almost a full month (see Table 1), our spectropolarimetric data set allows us to estimate, at the surfaces of both stars, the amount of latitudinal differential rotation shearing the brightness and/or magnetic maps. We achieve this in practice by

³ Note that these tilt and phase values refer to the visible magnetic pole, i.e. the negative and positive poles for V819 Tau and V830 Tau, respectively. Using the positive pole as reference (as in Gregory & Donati 2011), the tilt angle and phase of the dipole component of V819 Tau become 150° and 0.0.

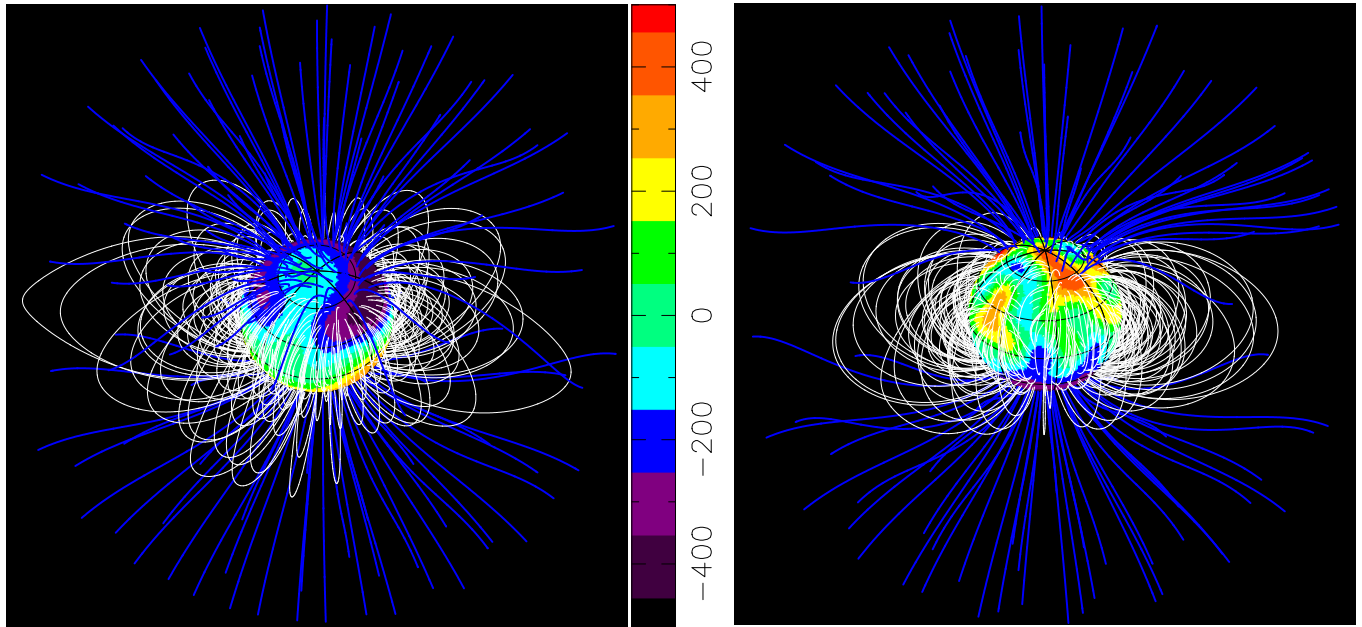


Figure 7. Potential field extrapolations of the magnetic topologies reconstructed for V819 Tau (left) and V830 Tau (right), as seen by an Earth-based observer at phases 0.0 and 0.15, respectively. Open and closed field lines are shown in blue and white, respectively, whereas colours at the stellar surface depict the local values (in G) of the radial field fluxes (as shown in the left panels of Fig. 6). The source surface at which the field becomes radial is (arbitrarily) set at a (realistic) distance of $4R_*$. This figure is best viewed in colour.

assuming that the rotation rate at the surface of the star is varying with latitude θ as $\Omega_{\text{eq}} - d\Omega \sin^2 \theta$ where Ω_{eq} is the rotation rate at the equator and $d\Omega$ the difference in rotation rate between the equator and the pole; in this context, one can measure differential rotation by finding out the pair of parameters Ω_{eq} and $d\Omega$ that produces the brightness/magnetic map with minimum information content. This procedure was successfully used in a large number of studies over the last 15 yr (e.g. Donati et al. 2000; Petit, Donati & Collier Cameron 2002; Donati, Collier Cameron & Petit 2003b), including on two wTTSs (Skelly et al. 2010; Donati et al. 2014). The reader is referred to these papers for more information on this technique.

Thanks to its high $v \sin i$ (see Table 3) and the good spatial resolution it offers, V830 Tau is an ideal target for estimating differential rotation; moreover, the presence of surface features (in both the brightness and magnetic maps) at various latitudes (see Figs 4 and 6) provides adequate information to obtain an accurate measurement. The χ_r^2 maps we obtain (as a function of both Ω_{eq} and $d\Omega$) are shown in Fig. 8; the one derived from our Stokes I data features a clear minimum at $\Omega_{\text{eq}} = 2.2950 \pm 0.0005 \text{ rad d}^{-1}$ and $d\Omega = 0.0124 \pm 0.0029 \text{ rad d}^{-1}$ (left panel of Fig. 8), whereas that inferred from our Stokes V data fully supports this estimate though with larger error bars ($\Omega_{\text{eq}} = 2.2950 \pm 0.0015 \text{ rad d}^{-1}$ and $d\Omega = 0.0152 \pm 0.0039 \text{ rad d}^{-1}$, right panel of Fig. 8). These values translate into rotation periods at the equator and pole of $2.738 \pm 0.001 \text{ d}$ and $2.753 \pm 0.004 \text{ d}$, respectively, fully consistent with the time-dependent periods of photometric fluctuations reported in the literature for V830 Tau (ranging from 2.738 to 2.747 d when considering only definite detections of photometric modulation; Grankin et al. 2008; Xiao et al. 2012) known to also probe surface differential rotation.

For V819 Tau, the spatial resolution of our maps is significantly coarser due to the lower $v \sin i$ (see Table 3); as a result, our Stokes

I data yield a χ_r^2 map showing no clear minimum, preventing us from estimating differential rotation. This is actually not very surprising given that the brightness map we recovered only includes two main features (one cool spot and one warm plage) more or less located at the same latitude (see Fig. 4); our data are thus lacking information about rotation periods and recurrence rates of profile distortions at different latitudes, without which differential rotation cannot be estimated (Petit et al. 2002). Our Stokes V data set yields a more favourable situation and allows us to obtain an estimate of surface differential rotation, equal to $\Omega_{\text{eq}} = 1.141 \pm 0.004 \text{ rad d}^{-1}$ and $d\Omega = 0.008 \pm 0.008 \text{ rad d}^{-1}$; unsurprisingly, the precision of this estimate is worse than for that of V830 Tau, reflecting the coarser spatial resolution achievable for this more slowly rotating star. These values translate into rotation periods at the equator and pole of $5.51 \pm 0.02 \text{ d}$ and $5.55 \pm 0.04 \text{ d}$, respectively. Our estimate is reasonably consistent with most photometric periods reported in the literature for V819 Tau (ranging from 5.50 to 5.61 d; Grankin et al. 2008; Xiao et al. 2012) except for one discrepant value (of 5.73 d, derived from data collected in 2004; Grankin et al. 2008). Looking back at the original photometric data, we find that this discrepant value actually corresponds to a low-accuracy estimate, reflecting the small number of measurements and the short observation time-span at this specific epoch.

Our results indicate that surface differential rotation is small on both V819 Tau and V830 Tau, much weaker in particular than that of the Sun (equal to $\approx 0.055 \text{ rad d}^{-1}$). Whereas we cannot exclude (from our data alone) that V819 Tau rotates as a solid body, the measurement we obtain for V830 Tau indicates that it truly rotates differentially and that the amount of surface shear that it suffers is ~ 4.4 times weaker than that of the Sun, with a typical time of $\approx 500 \text{ d}$ for the equator to lap the pole by one complete rotation cycle. Similarly small amounts of surface differential rotation were reported for the fully convective wTTSs LkCa 4 and V410 Tau

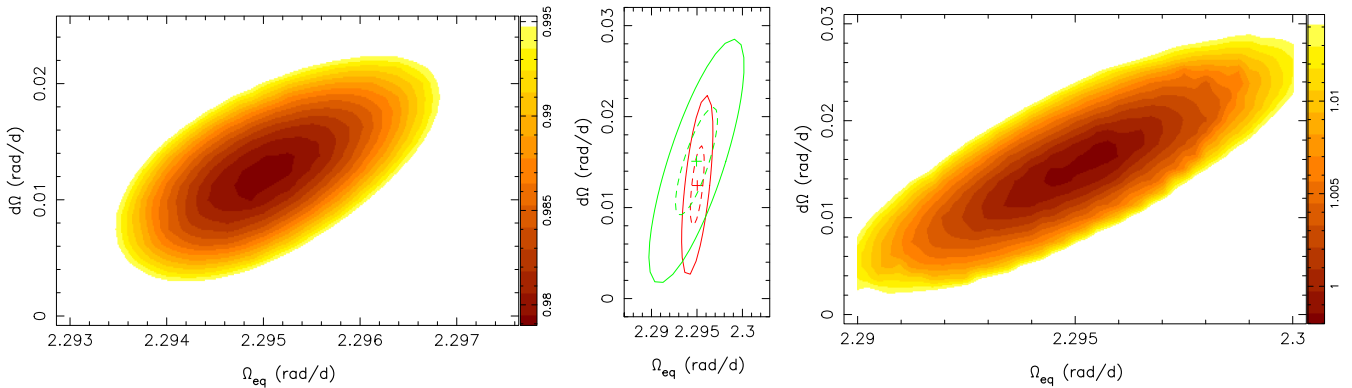


Figure 8. Variations of χ_r^2 as a function of Ω_{eq} and $d\Omega$, derived from the modelling of our Stokes *I* (left) and *V* (right) LSD profiles of V830 Tau at constant information content. In both cases, a clear and well-defined paraboloid is observed, with the outer colour contour tracing the 1.7 per cent increase in χ_r^2 (or equivalently a χ^2 increase of 11.8 for 705 fitted data points) that corresponds to a 3σ ellipse for both parameters as a pair. The middle plot emphasizes how well the confidence ellipses from both measurements overlap, with 1σ and 3σ ellipses (depicting respectively the 68.3 per cent and 99.73 per cent confidence levels) shown in full and dashed lines (in red and green for Stokes *I* and *V* data, respectively). This figure is best viewed in colour.

(Skelly et al. 2010; Donati et al. 2014) as well as for fully convective M dwarfs with presumably similar internal structures (Donati et al. 2006a; Morin et al. 2008a).

5 FILTERING THE ACTIVITY JITTER

As previously achieved for the wTTS LkCa 4 (Donati et al. 2014), we propose to exploit the results of surface imaging to filter out the activity jitter in the RV curves of V819 Tau and V830 Tau, and potentially reveal the presence of close-in giant planets (see Donati et al. 2014, for more details). Technically speaking, the goal is to use the brightness and magnetic maps reconstructed from our phase-resolved spectropolarimetric data sets with tomographic imaging to obtain an accurate description of the activity jitter that plagues the RV curves of our wTTSs; once RV curves are filtered out from the predicted activity jitter, one can look for periodic signals in the RV residuals that may probe the presence of hJs. Alternatively, we can *simultaneously* look for the presence of a planet while carrying out the imaging process, as recently proposed by Petit et al. (2015); in this case, the goal is to find out the optimal planet parameters, defined as those yielding the minimum amount of surface features for a given χ_r^2 fit to the data. Our initial simulations suggest that both techniques should allow one to detect potential hJs orbiting wTTSs (provided their periods are not too close to the stellar rotation period or its first harmonics) and to quantitatively investigate whether close-in giant planets are much more frequent around low-mass forming stars at the TTS stage than around mature MS stars. Ultimately, our goal is to assess the likelihood of the core accretion coupled to disc migration scenario in which giant protoplanets mostly form in outer accretion discs and then migrate inwards until they stop at a distance of a few hundredth of an au, e.g. once they enter the central magnetospheric gaps of cTTSs and no longer experience the inward torque from the disc.

The predicted activity jitter and filtered RV curves we derive for V819 Tau and V830 Tau are shown in Fig. 9. Note that for our study, RVs are estimated as first order moments of the Stokes *I* LSD profiles rather than through Gaussian fits, given the asymmetric and often irregular shapes of spectral lines (see Fig. 3, first and third panels). The filtering process is found to be quite efficient for V819 Tau, with RV residuals exhibiting an rms dispersion of only 0.033 km s^{-1} (for a full RV amplitude prior to the filtering process equal to $\simeq 1.6 \text{ km s}^{-1}$). The error bar on these measurements is

conservatively set to $\pm 0.04 \text{ km s}^{-1}$, reflecting mostly the intrinsic RV precision of ESPaDOnS (of 0.03 km s^{-1} rms; e.g. Moutou et al. 2007; Donati et al. 2008) and in a smaller part the intrinsic error of our filtering process (scaling up with the width of spectral lines, and estimated to be 0.02 km s^{-1} rms in the present case). This suggests that V819 Tau is unlikely to host a hJ with an orbital period in the range of what we can detect (i.e. not too close to the stellar rotation period or its first harmonics; see Donati et al. 2014), with a 3σ error bar on the semi-amplitude of the RV residuals equal to 0.07 km s^{-1} .

For V830 Tau, however, the RV curve exhibits significant scatter at several rotation phases (0.05, 0.35 and 0.65, see Fig. 9 right panel), causing the RV residuals after filtering to reach an rms dispersion of 0.104 km s^{-1} (i.e. about twice and thrice those reported for LkCa 4 and V819 Tau, respectively) and to exceed the value (of $\simeq 0.1 \text{ km s}^{-1}$) that potentially suggests the presence of a hJ (according to the results of our preliminary simulations; see Donati et al. 2014). The error bar on our individual RV residuals is set to $\pm 0.06 \text{ km s}^{-1}$, identical to those of LkCa 4 given the very similar widths of spectral lines and noise levels in our data. A closer look reveals that this enhanced dispersion is mainly caused by three points (at rotational cycles 3.654, 8.051 and 8.362; see Fig. 10) that depart (RV-wise) from the bulk of our observations. The temporal variations of the RV residuals are compatible with a sine wave with a period of either $4.0 \pm 0.2 \text{ d}$ or $6.0 \pm 0.2 \text{ d}$ and a semi-amplitude of $0.11 \pm 0.03 \text{ km s}^{-1}$. Bringing a χ^2 improvement of $\simeq 22$ for 4 degrees of freedom, this fit suggests that the RV residuals do indeed contain an RV signal, with a confidence level of 99.98 per cent. However, given the very limited sampling, we cannot firmly conclude that the RV signal we find is truly periodic as expected for a signal of planetary origin, the false alarm probability on the period detection (as derived from a Lomb–Scargle periodogram) being uncomfortably high (at about 35 per cent).

Experimenting the new technique proposed by Petit et al. (2015), whose advantage is to model both surface features and orbital parameters simultaneously, and thus to minimize crosstalk between the two and to maximize the precision of the whole process (at the expense of more computing time), we further confirm our finding, and conclude that the best tentative period of the signal in the RV residuals is $P_{\text{orb}} = 6.0 \pm 0.2 \text{ d}$, whereas the other option ($P_{\text{orb}} = 4.0 \pm 0.2 \text{ d}$) yields a worse fit to the data and is therefore less likely. The semi-amplitude of the residual RV signal that we derive in this case (taking into account differential rotation as

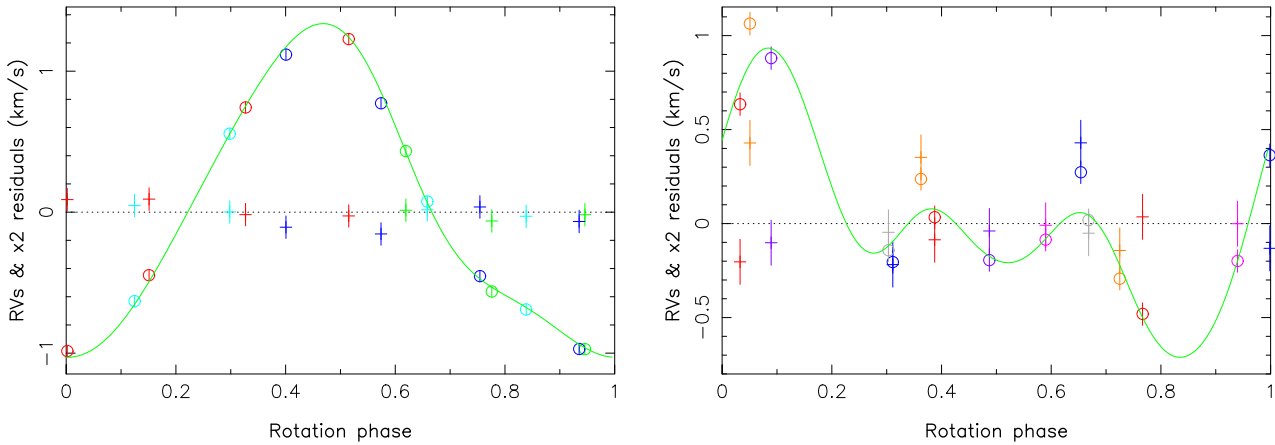


Figure 9. RV variations (in the stellar rest frame) of V819 Tau (left) and V830 Tau (right) as a function of rotation phase, as measured from our observations (open circles) and predicted by the tomographic maps of Fig. 4 (green line). RV residuals (expanded by a factor of 2 for clarity), are also shown (pluses) and exhibit an rms dispersion equal to 0.033 km s^{-1} for V819 Tau and 0.104 km s^{-1} for V830 Tau. Red, green, dark-blue, light-blue, magenta, grey, orange and purple symbols depict measurements secured at rotation cycles 0, 1, 3 (including phase 2.999), 4, 6, 7, 8 and 9, respectively (see Table 1 and Fig. 3). Given the asymmetric and often irregular shapes of Stokes I LSD profiles (see Fig. 3), RVs are estimated as first order moments of the Stokes I LSD profiles rather than through Gaussian fits. RV estimates and residuals are depicted with error bars of ± 0.04 and $\pm 0.06 \text{ km s}^{-1}$ for V819 Tau and V830 Tau, respectively. This figure is best viewed in colour.

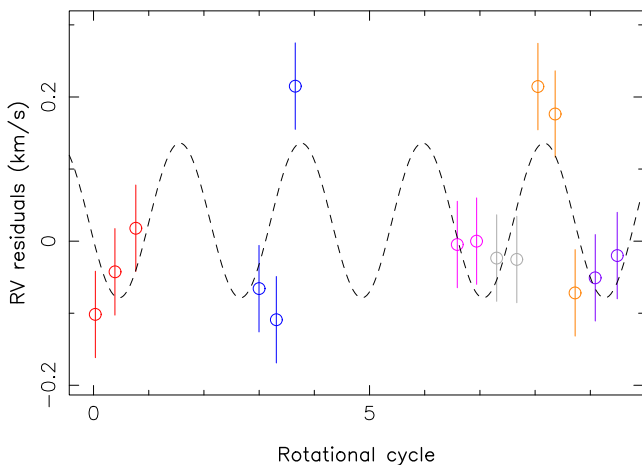


Figure 10. RV residuals of V830 Tau after applying our filtering process to the spectropolarimetric data set. A sine fit to the points with an orbital period of $2.2P_{\text{rot}} \simeq 6.0 \text{ d}$ and a velocity semi-amplitude of 0.11 km s^{-1} (dashed line) provides a good match to the observations (rms dispersion of 0.074 km s^{-1}), suggesting the possible presence of a giant planet orbiting close to V830 Tau. More observations are obviously needed to confirm or reject this preliminary result. The colour code used for this plot is the same as that of Fig. 9 (right panel). This figure is best viewed in colour.

measured in Section 4) is equal to $K = 0.13 \pm 0.02 \text{ km s}^{-1}$, consistent with, and presumably more accurate than, our initial estimate (of $0.11 \pm 0.03 \text{ km s}^{-1}$). A projection of the achieved χ^2 landscape in the K versus $P_{\text{orb}}/P_{\text{rot}}$ plane is shown in Fig. 11, demonstrating that our solution corresponds to a well-defined minimum. The associated improvement in χ^2 (with respect to a solution with no RV signal; see Fig. 8, left panel) is larger than 70 (for 4 degrees of freedom), implying a definite detection of a RV signal, with a very small false alarm probability ($< 10^{-10}$).⁴

⁴ This false alarm probability is smaller than that obtained by directly fitting RVs (equal to 0.02 per cent; see the previous paragraph), which confirms that line profiles contain more information than their first moments. Note

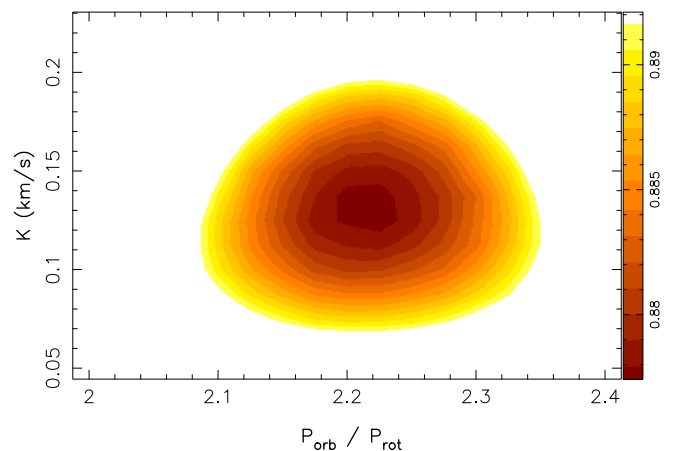


Figure 11. Variations of χ_r^2 as a function of $P_{\text{orb}}/P_{\text{rot}}$ and K , derived from the simultaneous modelling of the brightness features at the surface of V830 Tau and the orbital parameters of a close-in giant planet in circular orbit, following the new technique proposed by Petit et al. (2015). A clear minimum is obtained in the χ^2 landscape, whose projection in a K versus $P_{\text{orb}}/P_{\text{rot}}$ plane passing through the minimum is shown here, with the outer colour contour delimiting the 3σ (or 99.73 per cent) confidence region. This result suggests that a $\simeq 1.4 M_{\text{Jup}}$ planet may be orbiting around V830 Tau. This figure is best viewed in colour.

This residual RV signal, if confirmed, would be attributable to a $\simeq 1.4 M_{\text{Jup}}$ planet located at a distance of $\simeq 0.065 \text{ au}$ from V830 Tau, assuming a circular orbit with an orbital plane coinciding with the equatorial plane of V830 Tau. As expected, the impact of this residual RV signal on the reconstructed maps of Figs 4 and 6 is small, implying in particular that all conclusions of Section 4 are unaffected; this also applies to our differential rotation estimates of V830 Tau, which are not modified by more than a fraction of an error bar once the residual RV signal is removed, even for our most

that both tests only refer to the detection of an RV signal and not to its periodic nature (for which the false alarm probability is much larger).

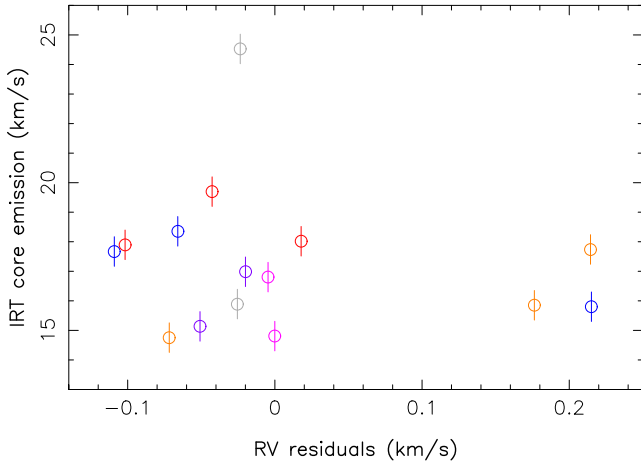


Figure 12. Equivalent width of the emission component in the core of Ca II IRT lines (in km s^{-1}) as a function of RV residuals for V830 Tau. No clear correlation is observed between both quantities, indicating that the observed RV residuals are likely not attributable to non-rotationally modulated activity. Note in particular that the three highest RV residuals all correspond to average levels of IRT core emission, whereas the highest level of IRT core emission (presumably probing a small flaring event) corresponds to an average RV. The colour code used for this plot is the same as that of Fig. 9 (right panel) and Fig. 10. This figure is best viewed in colour.

accurate measurement from Stokes I data. Similarly, the 3σ upper limit on the semi-amplitude of the RV residuals of V819 Tau (equal to 0.07 km s^{-1}) translates into a value of $\simeq 1.3 M_{\text{Jup}}$ for the mass of a planet orbiting at a distance of $\simeq 0.1 \text{ au}$ (and assuming again a circular orbit in the equatorial plane of the star).

We further checked that the temporal variations of the RV residuals we report here do not correlate with activity (e.g. estimated from the core emission in Ca II IRT lines; see Fig. 12), indicating that it cannot be attributed to, e.g. flaring events that may have damaged the profile shapes and corresponding RVs in our spectra. Similarly, we emphasize that these RV variations do occur on relatively short time-scales of only $\simeq 1 \text{ d}$ (e.g. from rotational cycle 3.311 to 3.654, and 8.362 to 8.725), suggesting that they are not attributable to some (slower) evolution of the brightness distribution at the surface of V830 Tau. Finally, we outline that RV variations are observed for V830 Tau only and not for V819 Tau, suggesting that they are likely not attributable to instrumental effects, which would equally plague spectra of both stars. We nevertheless stress that the detection of this RV signal currently relies on a small number of points, preventing us in particular to firmly assess its periodic nature; as a result, we cannot yet entirely exclude the option that it reflects some yet unclear systematics in our data.

6 SUMMARY AND DISCUSSION

In this paper, we report results from a new set of spectropolarimetric observations collected with ESPaDOnS at the CFHT on two wTTSs, V819 Tau and V830 Tau, from mid-2014 December to mid-2015 January, and complemented by contemporaneous photometric observations from the 1.25-m telescope at CrAO. As for our initial study on LkCa 4 (Donati et al. 2014), the monitoring of these two wTTSs was carried out in the framework of the international MaTYSSSE Large Programme.

From the analysis of our spectra, we find that both stars have similar atmospheric properties (photospheric temperature of $4250 \pm 50 \text{ K}$ and logarithmic gravity in cgs units of 3.9 ± 0.1),

suggesting that V819 Tau and V830 Tau are $1.00 \pm 0.05 M_{\odot}$ stars with respective radii of $1.8 \pm 0.2 R_{\odot}$ and $2.0 \pm 0.2 R_{\odot}$ viewed at inclination angles of 35° and 55° . With estimated ages of $\simeq 3.2$ and $\simeq 2.2 \text{ Myr}$, V819 Tau and V830 Tau are both close to the boundary at which $1 M_{\odot}$ stars start to build up a radiative core – with V819 Tau likely showing up as a slightly more evolved version than V830 Tau (note, however, the significant uncertainties on the ages; see Fig. 2). V819 Tau and V830 Tau also lie close to GQ Lup in the HR diagram (Donati et al. 2012), although they belong to a different star formation region.

With rotation periods of 5.53 and 2.74 d, V819 Tau and V830 Tau both spin significantly faster than typical cTTSs of similar mass (rotating in $\simeq 8 \text{ d}$, like GQ Lup); this suggests that both stars already entered a process of rapid spin-up after dissipating their accretion discs. We note that V819 Tau rotates slower despite being older than V830 Tau; this suggests that V830 Tau spent more time spinning up since the dissipation of its (presumably less massive) accretion disc than V819 Tau and thus that both stars had slightly different accretion histories. Although speculative, this scenario would also be consistent with the fact that V819 Tau exhibits traces of leftover dust from the original disc whereas V830 Tau shows none (e.g. Cieza et al. 2013).

Applying our tomographic imaging code (adapted to the case of wTTSs; see Donati et al. 2014) to our new data set, we derived the surface brightness maps and magnetic topologies of both stars. Cool spots and warm plages are found to be present at the surfaces of both V819 Tau and V830 Tau, but with a smaller contrast than the brightness features previously mapped on LkCa 4. This is especially true for V830 Tau despite the fact that it rotates about twice as fast as V819 Tau and was thus expected to have built up, e.g. a large high-contrast polar spot like that of V410 Tau or many other similar rapidly rotating young low-mass stars. This may suggest that V830 Tau was in a lower-than-usual state of activity at the time of our observations.⁵ As for LkCa 4, we find that the brightness maps reconstructed for both stars are in very good agreement with our contemporaneous photometric observations.

The large-scale magnetic fields we reconstruct for V819 Tau and V830 Tau are similar, both in terms of intensities and topologies, and are found to be 80–90 per cent poloidal; the poloidal component is dominated (especially at some distance from the stars) by a dipolar term of polar strength 350–400 G inclined at $\simeq 30^{\circ}$ to the rotation axis whereas the octupolar term is weaker than the dipolar one. We stress that neither star shows a conspicuous ring of strong toroidal fields as that reported for LkCa 4 (Donati et al. 2014). The fact that both stars have similar fields despite their $\times 2$ ratio in rotation rates does not come as a surprise given that dynamo action is expected to be largely saturated for these active wTTSs. More unexpected is that these fields are much weaker than that of GQ Lup (typically hosting a 1 kG dipole and a 2 kG octupole; see Donati et al. 2012) despite their proximity in the HR diagram. Although a topological difference (the afore-mentioned toroidal component) was also reported between the large-scale fields of LkCa 4 on the one hand, and those of AA Tau and BP Tau on the other hand (located near LkCa 4 in the HR diagram), their poloidal components were similar in terms of both strengths and topologies.

⁵ Although long-term photometric monitoring shows that V830 Tau was indeed exhibiting lower-than-usual photometric variability at the time of our observations (Grankin et al. 2008), previous studies (e.g. Skelly et al. 2010) demonstrated that overall activity of low-mass stars, and more specifically the size of their polar spots, is not necessarily correlated with the amplitude of photometric variability.

It is obviously still too early to assess whether magnetic topologies of cTTSs and wTTSs are similar or dissimilar in the regions of the HR diagram where both can coexist. Our observations so far suggest that they differ either in their poloidal components (this new result compared to Donati et al. 2012) or in their toroidal components (the case of LkCa 4 versus AA Tau and BP Tau; Donati et al. 2014). The reason for this difference, if confirmed, will need clarification and is potentially complex, acting apparently on both field strengths and topologies. Having the potential power to significantly affect atmospheric properties of cTTSs (and thus their location in the HR diagram; Baraffe, Chabrier & Gallardo 2009) and/or to alter their large-scale topologies (by stressing them through disc/star mass transfers and/or by modifying the underlying dynamo processes), accretion can participate in this process. However, the similarity between magnetic topologies of cTTSs and M dwarfs with comparable internal structures (Gregory et al. 2012; Donati et al. 2014) suggests the opposite. The temporal variability inherent to dynamo processes is another option; repeated observations of magnetic fields of cTTSs and wTTSs so far (e.g. Donati et al. 2011, 2012, 2013), however, indicate that the amplitude of such variations with time is smaller than the one we report here. Clearly, we will have to wait for more MaTYSSSE studies in the line of our two first ones to progress on this issue.

Our data also clearly demonstrate that brightness maps and/or magnetic topologies of V830 Tau and V819 Tau experience very little latitudinal shear in a time-scale of $\simeq 1$ month. In the particular case of V830 Tau, whose high $v \sin i$ yields the best spatial resolution at the surface of the star, we estimated the amount of differential rotation at a high-enough precision to claim that it is different from zero (i.e. not rotating as a solid body) at a confidence level larger than 99.99 per cent, and smaller than that of the Sun by typically $4.4\times$ (with a time for the equator to lap the pole by one complete rotation cycle equal to $\simeq 500$ d, as opposed to $\simeq 110$ d for the Sun); moreover, the differential rotation estimates we derived from the brightness and magnetic maps of V830 Tau, as well as that obtained from the magnetic maps of V819 Tau, all agree together within the error bars. Our results also agree with the few existing estimates of differential rotation at the surfaces of similar low-mass wTTSs (Skelly et al. 2010; Donati et al. 2014).

As for LkCa 4, we finally exploited the results of our tomographic imaging of V819 Tau and V830 Tau to model and filter-out the activity jitter in the RV curves of both stars. For V819 Tau, we find that the activity jitter is filtered down to an rms RV precision of 0.033 km s^{-1} , further demonstrating that our technique works well and performs as expected from our initial set of simulations (Donati et al. 2014). This leaves little hope to find potential hJs around V819 Tau using our technique despite its exhibiting traces of dust, possibly indicating planet formation, which made it an obvious target for our programme. For V830 Tau, however, the RV residuals after filtering exhibit a significantly larger rms dispersion of 0.104 km s^{-1} caused by variability on time-scales of a few days; we find that this dispersion excess is unlikely to relate to intrinsic (e.g. flaring-like) activity (since RV residuals do not correlate with activity proxies) nor to instrumental effects (that would impact data from both stars). One option is that the observed excess dispersion is caused by a $\simeq 1.4 M_{\text{Jup}}$ giant planet orbiting V830 Tau in 6.0 ± 0.2 d at a distance of 0.065 au from the central star. Although our detection of a residual RV signal is in principle reliable with a high confidence level (given the excess RV dispersion it generates), we nevertheless stress that it relies on a small number of points, preventing us to firmly assess its periodic nature; as a result, we cannot exclude the option that it reflects yet unclear

systematics in our data rather than the true presence of a giant planet.

Additional data on V830 Tau with a denser sampling of both rotational and orbital cycles are required to obtain a definite and independent validation of this first promising result. If confirmed, our result would suggest that hJs are likely much more frequent around wTTSs than around MS stars, to ensure that one is detected in a randomly selected sample of only a few stars. The ongoing MaTYSSSE observations will enable to further clarify this key issue for our understanding of the formation of planetary systems. Thanks to its much higher sensitivity to wTTSs and enhanced RV precision, SPIRou, the nIR spectropolarimeter/high-precision velocimeter currently in construction for CFHT (first light planned in 2017), will tremendously boost this research programme in the near future.

ACKNOWLEDGEMENTS

This paper is based on observations obtained at the CFHT, operated by the National Research Council of Canada, the Institut National des Sciences de l'Univers of the Centre National de la Recherche Scientifique (INSU/CNRS) of France and the University of Hawaii. We thank the CFHT QSO team for its great work and effort at collecting the high-quality MaTYSSSE data presented in this paper. MaTYSSSE is an international collaborative research programme involving experts from more than 10 different countries (France, Canada, Brazil, Taiwan, UK, Russia, Chile, USA, Switzerland, Portugal, China and Italy). We also warmly thank the IDEX initiative at Université Fédérale Toulouse Midi-Pyrénées (UFTMiP) for funding the STEPS collaboration program between IRAP/OMP and ESO and for allocating a 'Chaire d'Attractivité' to GAJH, allowing her to regularly visit Toulouse to work on MaTYSSSE data. We acknowledge funding from the LabEx OSUG@2020 that allowed purchasing the ProLine PL230 CCD imaging system installed on the 1.25-m telescope at CrAO. SGG acknowledges support from the Science & Technology Facilities Council (STFC) via an Ernest Rutherford Fellowship [ST/J003255/1]. SPHA acknowledges financial support from CNPq, CAPES and Fapemig. AAV acknowledges support from the Swiss National Science Foundation (SNSF) via the allocation of an Ambizione Fellowship.

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