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## THE NEXT GENERATION VIRGO CLUSTER SURVEY. VII. THE INTRINSIC SHAPES OF LOW-LUMINOSITY GALAXIES IN THE CORE OF THE VIRGO CLUSTER, AND A COMPARISON WITH THE LOCAL GROUP

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### ABSTRACT

We investigate the intrinsic shapes of low-luminosity galaxies in the central 300 kpc of the Virgo cluster using deep imaging obtained as part of the Next Generation Virgo Cluster Survey (NGVS). We build a sample of nearly 300 red-sequence cluster members in the yet unexplored  $-14 < M_g < -8$  magnitude range, and we measure their apparent axis ratios,  $q$ , through Sérsic fits to their two-dimensional light distribution—which is well described by a constant ellipticity parameter. The resulting distribution of apparent axis ratios is then fit by families of triaxial models with normally-distributed intrinsic ellipticities,  $E = 1 - C/A$ , and triaxialities,  $T = (A^2 - B^2)/(A^2 - C^2)$ . We develop a Bayesian framework to explore the posterior distribution of the model parameters, which allows us to work directly on discrete data, and to account for individual, surface brightness-dependent axis ratio uncertainties. For this population we infer a mean intrinsic ellipticity  $\bar{E} = 0.43_{-0.02}^{+0.02}$ , and a mean triaxiality  $\bar{T} = 0.16_{-0.06}^{+0.07}$ . This implies that faint Virgo galaxies are best described as a family of thick, nearly oblate spheroids with mean intrinsic axis ratios  $1 : 0.94 : 0.57$ . The core of Virgo lacks highly elongated low-luminosity galaxies, with 95 per cent of the population having  $q > 0.45$ . We additionally attempt a study of the intrinsic shapes of Local Group (LG) satellites of similar luminosities. For the LG population we infer a slightly larger mean intrinsic ellipticity  $\bar{E} = 0.51_{-0.06}^{+0.07}$ , and the paucity of objects with round apparent shapes translates into more triaxial mean shapes,  $1 : 0.76 : 0.49$ . Numerical studies that follow the tidal evolution of satellites within LG-sized halos are in good agreement with the inferred shape distributions, but the mismatch for faint galaxies in Virgo highlights the need for more adequate simulations of this population in the cluster environment. We finally compare the intrinsic shapes of NGVS low-mass galaxies with samples of more massive quiescent systems, and with field, star-forming galaxies of similar luminosities. We find that the intrinsic flattening in this low-luminosity regime is almost independent of the environment in which the galaxy resides—but there is a hint that objects may be slightly rounder in denser environments. The comparable flattening distributions of low-luminosity galaxies that have experienced very different degrees of environmental effects suggests that internal processes are the main drivers of galaxy structure at low masses—with external mechanisms playing a secondary role.

*Subject headings:* galaxies: clusters: individual (Virgo) – Local Group – galaxies: dwarf – galaxies: photometry – galaxies: fundamental parameters – galaxies: structure

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## 1. INTRODUCTION

The faint end of the satellite luminosity function is dominated by systems with smooth appearance, low surface brightness, and little gas content or recent star formation activity—the so-called dwarf spheroidal (dSph) galaxies (Mateo 1998; van den Bergh 2000; Grebel et al. 2003; Tolstoy et al. 2009, and references therein). In the Local Group, their luminosities and stellar velocity dispersions are similar to those of globular clusters, but their much larger half-light radii imply substantial total masses, making these faint galaxies the most dark matter dominated systems known in the Universe (e.g., Walker et al. 2009).

Two observed properties of this population of faint satellites are known to place strong constraints on  $\Lambda$ CDM models of galaxy formation. First, their number is short by one to two orders of magnitude compared to the expectations from the mass function of cold dark matter halos (Klypin et al. 1999; Moore et al. 1999; but see Hargis et al. 2014). Attempts to resolve this discrepancy traditionally invoke baryonic mechanisms that lower the efficiency of, or even suppress, star formation at these mass scales—including high gas cooling times, reionization, stellar feedback, and early environmental effects (White & Rees 1978; Dekel & Silk 1986; Bullock et al. 2000; Benson et al. 2003; Mayer et al. 2007; Okamoto et al. 2008). In passing, we note that the abundance of the gas-rich, star-forming counterparts which exist in the same luminosity regime in the field also seems to be lower than what is predicted by  $\Lambda$ CDM (Klypin et al. 2015). Obviously, environmental effects can not be invoked in this case. Second, the inferred central dark matter densities in these systems favor constant density cores instead of the steeper density profiles found in dark matter simulations (Walker & Peñarrubia 2011). Hydrodynamic simulations suggest that the removal of these dark matter cusps occurs through energetic outflows from supernovae, a mechanism that requires frequent and centrally concentrated bursts of star formation (Governato et al. 2010; Pontzen & Governato 2012; Teyssier et al. 2013).

Peñarrubia et al. (2012) point out that these two effects place competing requirements for the star formation efficiency in low-mass halos, creating tension in CDM models of galaxy formation. This friction can be alleviated, however, by a combination of baryonic physics and (enhanced) tidal stripping that bring both the inner mass profiles and the abundance of satellites in simulations in agreement with the observations (Zolotov et al. 2012; Brooks et al. 2013). Indeed, low-mass quiescent satellites tend to be found in the vicinity of luminous ( $L \gtrsim L^*$ ) central galaxies (Geha et al. 2012; Sánchez-Janssen et al. 2013), and dynamical processes have long been proposed to influence their evolution (Peñarrubia et al. 2008; Kazantzidis et al. 2011).

In order to put constraints on the plethora of internal and external mechanisms potentially shaping the properties of faint satellites, it is mandatory to robustly characterize these systems observationally. In this sense, it is only natural to start with the most accessible observables, i.e., their shapes and structural parameters. Importantly, the relatively shallow potential wells in which these galaxies reside ensure that both photoionization and stellar feedback processes, as well as environmental

effects, leave recognizable imprints on their stellar structure. For instance, Kaufmann et al. (2007) show that baryonic processes that act to pressurize gas within low-mass halos naturally lead to the formation of thicker, puffier galaxies towards lower luminosities as a result of their progressively diminishing angular momentum support. In particular, feedback from bursty star formation can result in strong potential fluctuations caused by gas outflows and inflows. The ensuing heating of stellar orbits can drive significant structural variation, including outward migration and thickening of the stellar body (Teyssier et al. 2013; El-Badry et al. 2015). Other sources of internal heating, such as dynamical interactions with dark subhalos, may play a role as well (Helmi et al. 2012; Starkenburg & Helmi 2015). Alternatively, tidal effects have the potential to significantly modify the structural (and dynamical) properties of disk systems, creating pressure-supported triaxial objects that in many aspects resemble the satellites of cluster- and group-sized halos (Mastropietro et al. 2005; Smith et al. 2010; Kazantzidis et al. 2011; Lokas et al. 2012; Bialas et al. 2015). Gas stripping mechanisms have a smaller direct impact on the stellar structure of low-mass systems, except for extreme cases where the gas component represents a large fraction of the total mass budget (see Smith et al. 2013a). They can nevertheless contribute to the thickening of the stellar body on timescales of a few Gyr by cutting the supply of young stars with low velocity dispersion, followed by subsequent diffusion through two-body relaxation.

Until recently, the intrinsic shapes of faint quiescent galaxies have not been studied extensively. This stems from the fact that these are intrinsically faint objects, and the great majority of known and well-characterized dSphs are Local Group satellites—and the number statistics there are low. For example, the complete sample of Milky Way (MW) dSphs studied by Martin et al. (2008) includes only 21 galaxies, 11 of which are fainter than  $M_V = -7.5$  mag. Interestingly, while the classical MW satellites show small apparent ellipticities, the faint population display more elongated shapes. This immediately prompts the suspicion that tidal interactions with the MW are behind the origin of this highly elongated galaxies. Along these lines, Lokas et al. (2012) find good general agreement between the apparent shapes of MW satellites and simulations of tidally stirred disks orbiting an MW-sized halo. Recently, Salomon et al. (2015) have inferred the intrinsic ellipticity of M31 satellites under the assumption that their visible components have prolate shapes, and after exploring different possibilities for the orientation of the major axis. They find that the population has a mean intrinsic axis ratio  $\approx 0.5$ , but several satellites are consistent with being intrinsically more elongated. Because the orientation of individual galaxies in the sky is hard to obtain, intrinsic shapes are usually derived on statistical grounds using relatively large samples. Low number statistics still make this task challenging in the Local Group (but again, see Salomon et al. 2015). It is however now possible to study large numbers of faint, quiescent galaxies in the nearby Virgo cluster thanks to the Next Generation Virgo Cluster Survey (NGVS).

The NGVS (Ferrarese et al. 2012, F12 hereafter) is a Canada-France-Hawaii Telescope (CFHT) optical imag-

ing survey of the Virgo cluster covering  $104 \text{ deg}^2$ , from its core to the virial radius. The MegaCam deep imaging delivers a  $2\sigma$  surface brightness limit  $\mu_g \approx 29 \text{ mag arcsec}^{-2}$ , with excellent image quality (median FWHM = 0.54 arcsec in the  $i$ -band). This has allowed us to detect and characterize, in the core region alone ( $R_p \lesssim 300 \text{ kpc}$ ), hundreds of previously unknown low-mass galaxies down to the luminosity limits of the classical MW satellites,  $M_V \sim -8 \text{ mag}$ . This constitutes a major breakthrough, not only because of the order(s) of magnitude increase in the sample size, but also because Virgo represents a completely different environment from the Local Group. Here we use the observed apparent axis ratios of a sample of nearly 300 low mass galaxies in the core of the Virgo cluster to infer the distribution of their intrinsic shapes. We will show that these low-mass Virgo satellites are a family of thick, nearly oblate spheroids.

This paper is organized as follows. In Section 2 we present the dataset, and the detection and structural characterization of faint Virgo galaxies. We then describe the Bayesian framework developed to infer their intrinsic shapes in Section 3, followed by the main results of our analysis in Section 4. In Section 5 we discuss our results, and compare them with previous efforts in the literature. Finally, in Section 6 we summarize our findings.

Throughout this work we adopt the convention that apparent (i.e., observed) quantities are represented with lower case symbols, whereas upper case symbols are used for intrinsic parameters. Thus, a triaxial galaxy with intrinsic semi axes  $A \geq B \geq C$  has an apparent ellipticity ( $\epsilon$ ) and an apparent axis ratio ( $q$ ) related through  $\epsilon = 1 - q$ . We furthermore use a common distance modulus ( $m - M$ ) = 31.09 mag for all candidate Virgo members, corresponding to the mean distance of  $D = 16.5 \text{ Mpc}$  to the Virgo cluster derived through the surface brightness fluctuations method (Mei et al. 2007; Blakeslee et al. 2009).

## 2. DETECTION AND CHARACTERIZATION OF LOW-MASS QUIESCENT GALAXIES IN VIRGO

### 2.1. Sample selection in the NGVS

In this work we use data from the NGVS. Here we only provide a brief overview of the most relevant characteristics of the survey, but the interested reader is referred to F12 for a more detailed description of the survey, datasets and science goals. The NGVS is a CFHT/MegaCam imaging survey of the Virgo cluster covering  $104 \text{ deg}^2$  in the  $u^*giz$  bandpasses, supplemented with  $K_s$ -band imaging for a fraction of the area (Muñoz et al. 2014), as well as a series of dedicated spectroscopic follow-ups. The central  $4 \text{ deg}^2$  around M87 ( $R \lesssim 300 \text{ kpc}$ ) were observed in a Pilot Programme (PP) that included the  $r$ -band as well, and all objects in this study come from this core region. By design, the survey has an excellent image quality, with a median seeing FWHM  $\approx 0.8 \text{ arcsec}$  in the  $g$ -band (corresponding to  $\sim 65 \text{ pc}$  at the distance of Virgo). The *Elxir-LSB* reduction pipeline (Ferrarese et al. 2012; Duc et al. 2015) delivers a superb flattening of the sky background, with peak-to-peak residuals of  $\sim 0.2$  per cent that result in a  $2\sigma$  surface brightness limit  $\mu_g \approx 29 \text{ mag arcsec}^{-2}$  for the final stacked frames. The NGVS is therefore perfectly positioned to carry out, for the first time, an investiga-

tion of the properties of low-mass satellites in the Virgo cluster. The identification and characterization of Virgo cluster members will be discussed in detail in an upcoming paper of this series (Ferrarese et al., submitted); here we summarize the most relevant aspects.

The detection of extended, low surface brightness Virgo galaxies in a field contaminated with a far larger number of foreground stars and background objects is a challenge; conventional codes, such as SExtractor (Bertin & Arnouts 1996), tend to regard low surface brightness objects as belonging to compact and/or brighter contaminants. To circumvent the problem, a ring median filter (Secker 1995) was first applied to each  $g$ -band image (the NGVS images with the highest signal-to-noise ratio), with radius adjusted to suppress all unresolved sources (stars and globular clusters), as well as compact background galaxies. A specific optimization of SExtractor was then run on the median-kernel smoothed images, thus enabling the identification of low surface brightness sources that are potential Virgo members.

Our low-mass galaxy sample explicitly excludes the class of so-called ultra-compact dwarfs (UCDs; Hilker et al. 1999; Drinkwater et al. 2000). These objects have sizes and luminosities bridging the main locii occupied by low-mass galaxies and globular clusters in structural scaling relations. As a result, the nature of UCDs is still hotly debated in the literature, and they probably encompass objects of both galactic and star cluster origin (e.g., Håegem et al. 2005; Mieske et al. 2008; Chiboucas et al. 2011; Penny et al. 2014; Norris et al. 2014; Seth et al. 2014). We refer the interested reader to Durrell et al. (2014) for an overview of the spatial distribution of bright compact stellar systems in Virgo; to Zhang et al. (2015), where we have published a kinematical analysis of  $\sim 100$  spectroscopically confirmed UCDs in the same Virgo core area we study here; and to Liu et al. (2015) for an investigation of their photometric properties.

The resulting galaxy catalogue does of course not discriminate between Virgo members and contaminants, and therefore the next challenge is to define membership criteria. These are based on the location of each galaxy in a multi-parameter space defined by a combination of i) galaxy structural parameters, specifically size and surface brightness, measured in appropriately masked original images using GALFIT (Peng et al. 2002); ii) photometric redshifts,<sup>1</sup> based on  $u^*, g, i, z$ , and, when available,  $r$ -band photometry; and iii) an index measuring the strength of residual structures in images created by subtracting from each galaxy the best fitting GALFIT model. The exact combination of axes in this space was selected to allow for maximum separation of known Virgo and background sources. The former comprise spectroscopically confirmed (mostly) VCC galaxies, while the latter are identified, using the same procedure described above, in four control fields located three virial radii away from M87, and presumed to be devoid of cluster members.

The reliability of the procedure was tested in two independent ways. First, 36,500 artificial galaxies were

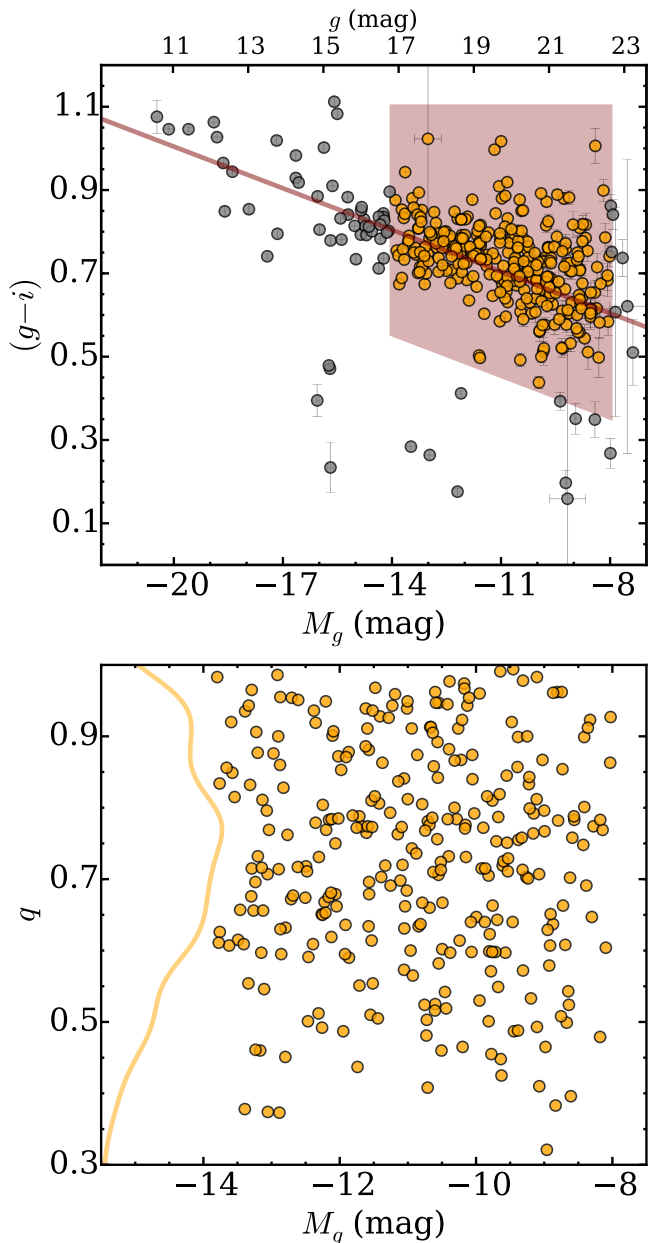
<sup>1</sup> We note here that photometric redshifts are mainly useful to identify and discard obvious background systems rather than selecting Virgo members, as their typical accuracy at low redshift,  $\delta z/(1+z) \sim 0.05$ , is not good enough to establish secure cluster membership (e.g., Raichoor et al. 2014).

injected in the frames spanning the full range of luminosity and structural parameters expected for genuine Virgo galaxies, and then processed as described above. With few exceptions (for instance galaxies that land in the immediate vicinity of bright saturated stars or near the cores of high surface brightness galaxies), all galaxies with surface brightness high enough to be visible in the frames were indeed recovered. A detailed discussion of the completeness of the data and the biases in the recovered parameters is included in Ferrarese et al. (submitted). Second, four NGVS team members (LF, PAD, PD, and FB) independently inspected the central four square degrees of the PP region and identified, by eye, all objects that appeared to be bona-fide Virgo members; again with few exceptions all such galaxies were detected and correctly identified as bona-fide Virgo members by the code. We stress here that the proximity of the Virgo cluster, together with the remarkable tightness of galaxy scaling relations even at these faint luminosities, are key in providing a clean separation between cluster members and background galaxies even in the absence of spectroscopic information (see also Michard & Andreon 2008; Rines & Geller 2008).

Figure 1 (top) presents the  $(g-i)$  color-magnitude diagram of all 404 candidate Virgo galaxies in the PP area (except M87). As expected given that we cover the inner regions of the cluster, the great majority of objects fall on a well-defined red sequence all the way down to the faintest luminosities probed (see also Sánchez-Janssen et al. 2008; Boselli & Gavazzi 2014). In this work we focus on the shapes of faint, low surface brightness, quiescent galaxies, analogous to the dSph satellites in the LG. It is now well established that this label encompasses objects with quite varied star formation histories (e.g., Weisz et al. 2014). Throughout this work we keep this terminology when referring to LG satellites for historic reasons, but we will avoid it as much as possible when referring to faint, quiescent Virgo galaxies. Because detailed star formation histories are not available yet for Virgo galaxies of these low masses, we construct our sample of quiescent objects from galaxies with colors consistent with being red-sequence cluster members (see the shaded region in the top panel of Fig. 1). The lower boundary is set at 0.25 mag below the best-fit color-magnitude relation, whereas  $(g-i) = 1.1$  defines the red end of our selection function. Furthermore, we only select objects in the  $-14 \leq M_g \leq -8$  mag range.<sup>2</sup> With this luminosity cut we guarantee our galaxies to be genuinely faint systems; we match the luminosity range of other low-luminosity samples in the LG and in the Local Universe that will serve as comparison; and we probe a yet unexplored galaxy mass range in Virgo—or any other cluster, for that matter.

We also explored a visual morphological classification with the aid of composite  $giz$  color images and unsharp-masked  $g$ -band frames. Faint quiescent systems were defined as objects with no evident star formation activity and smooth, elliptical appearances. We found that both the color and morphological classifications produced very similar samples, and therefore our results do not strongly

<sup>2</sup> This roughly corresponds to stellar masses  $3 \times 10^5 \lesssim M_*/M_\odot \lesssim 7 \times 10^7$ , for a characteristic stellar mass-to-light ratio  $M_*/L_V = 1.6 M_\odot/L_\odot$  (Woo et al. 2008).

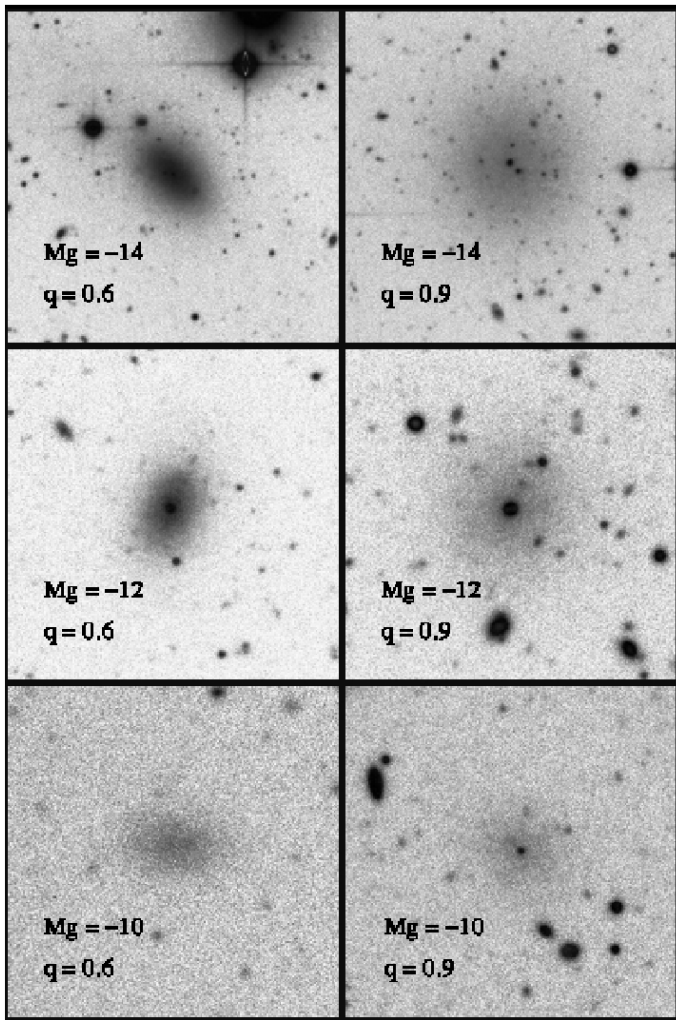


**Figure 1.** *Top:*  $(g-i)$  color-magnitude diagram of galaxies in the core of the Virgo cluster. Low-luminosity Virgo candidates are selected as galaxies having absolute magnitudes in the  $-14 \leq M_g \leq -8$  mag range, and colors consistent with being red-sequence cluster members—as indicated by the shaded region. *Bottom:* apparent axis ratios for the low-luminosity sample as a function of absolute magnitude. The solid line is a simple gaussian kde estimate of the  $q$  distribution. Note the absence of very elongated galaxies, and the lack of dependence on luminosity.

depend on the exact definition of what is a quiescent low-mass galaxy. In the following analysis we use the color selection simply because it results in a more quantitative and reproducible selection function.

## 2.2. Measurements of apparent axis ratios

Structural parameters for low-luminosity galaxies in the NGVS are obtained from both surface brightness profiles using the IRAF task ELLIPSE (Jedrzejewski 1987), and two-dimensional model fits using GALFIT. The derived structural parameters include, but are not limited



**Figure 2.** NGVS negative postage stamps of six representative faint Virgo galaxies. Rows correspond to galaxies with similar absolute magnitudes, while systems with similar apparent axis ratios are shown along columns. The left column displays objects with significant apparent elongation ( $q \approx 0.6$ ), while nearly round ( $q \approx 0.9$ ) objects are shown on the right column.

to, the effective radius and surface brightness, the Sérsic shape parameter, the position angle, and the apparent axis ratio  $q$ —which is the most relevant quantity for this study.<sup>3</sup> Here we will only use structural parameters derived in the  $g$ -band, as it is the band with the highest signal-to-noise ratio in the NGVS. We have verified that our results do not change if we use either the  $r$ - or the  $i$ -band data. This is expected if these faint systems, like the dSphs in the LG, have relatively uniform stellar populations, and little obscuration caused by dust.

In both approaches the galaxy body is modeled with a Sérsic (1968) function, but only in the case of the isophotal analysis with ELLIPSE is the ellipticity allowed to vary with radius—GALFIT uses constant axis ratio models. If necessary, a central nuclear component is included in the fits. With typical half-light radii  $r_h < 0.3$  arcsec at the distance of Virgo (Côté et al. 2006; Turner et al. 2012), most of these nuclei remain unresolved in the NGVS images. They are simply modelled as PSF components (see

<sup>3</sup> In this work we use the mean axis ratio derived by ELLIPSE in the  $1 \leq r/r_e \leq 2$  interval as a point estimate for  $q$ .

also den Brok et al. 2014). Finally, all the parameters derived from GALFIT fits are seeing-deconvolved, while ELLIPSE measures seeing-convolved axis ratios. Once again, this should not impact the global structural parameters of objects that have shallow surface brightness profiles and are well resolved. For reference, the smallest effective radius for an  $M_g \approx -8$  Virgo galaxy in our sample is  $r_e \approx 2.5$  arcsec ( $\sim 200$  pc), which is more than a factor three larger than the  $g$ -band FWHM.

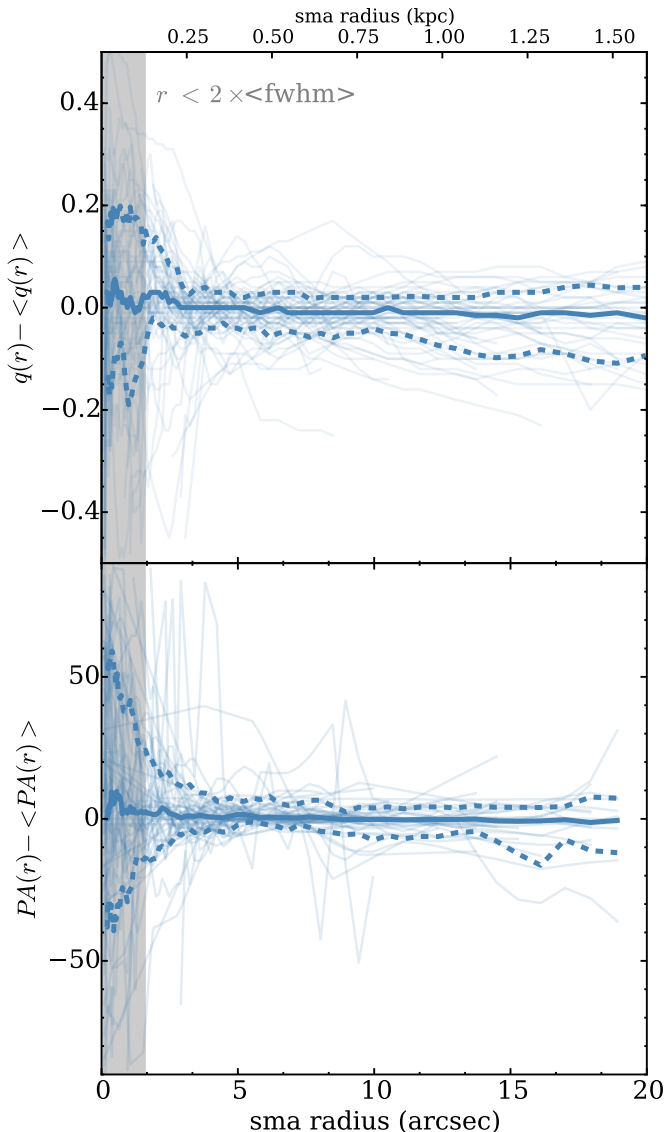
In this work we use the GALFIT axis ratios for a practical reason. For a non-negligible fraction of the faintest galaxies, the ELLIPSE fits do not converge when all parameters are allowed to vary due to a combination of small apparent ellipticity and very low central surface brightness. For these objects the fits were carried out with the axis ratio fixed at  $q = 0.95$ . Using the ELLIPSE parameters would create an artificial peak in the distribution of observed axis ratios at this value, and would prevent us from modeling it appropriately (see Sect. 3).

The observed distribution of GALFIT axis ratios for our faint Virgo galaxies is plotted as a function of  $g$ -band absolute magnitude in the bottom panel of Figure 1. It is interesting to note the paucity of low-luminosity systems in the core of Virgo displaying highly elongated shapes: 95 per cent of the population have  $q \gtrsim 0.45$ . This is an important result considering we are only sampling objects with projected distances  $\lesssim 300$  kpc from M87, that is, the highest density regions of the cluster where tidal forces are maximized. Additionally, and at least in this mass regime and cluster region, there does not seem to be any dependence of the apparent shapes  $q$  on luminosity or clustercentric distance. This guarantees that we can use the entire sample to infer the intrinsic shape of the population. Figure 2 shows negative postage stamps of six representative galaxies in the PP region, illustrating the range of luminosities and apparent shapes of this faint population, and the quality of the NGVS imaging.

### 2.2.1. Low-luminosity galaxies in Virgo have flat ellipticity profiles

One potential limitation of the GALFIT surface-brightness-weighted  $q$  estimates is that they do not necessarily represent the shape of the galaxy at any particular radius. This could introduce a bias if, for example, the ellipticity profile presented significant radial variation, or if the amplitude of that variation was luminosity-dependent. We can rule out this is the case for our sample of NGVS low-mass objects from an analysis of the quality of the GALFIT fits. Sérsic models with constant ellipticity provide excellent fits to this population, with a median  $\chi^2_\nu = 1.00$ , and a dispersion  $\sigma_{\chi^2_\nu} = 0.11$ . Moreover, fit residuals do not show any radial structure, thus ruling out strong isophotal twisting as well.

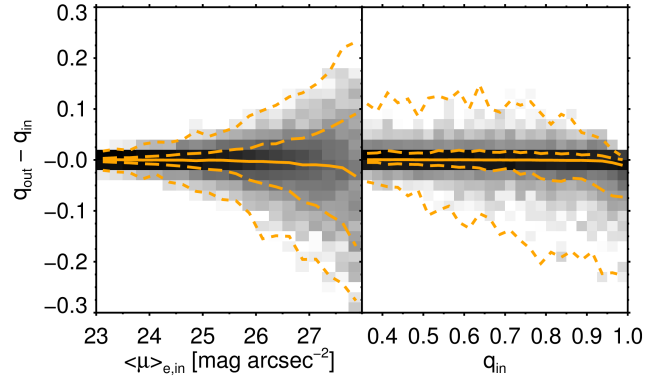
We also use the ELLIPSE fits to study the radial ellipticity profiles in our sample. Thin lines in Fig. 3 show, for objects having  $\langle \mu \rangle_e < 27$  mag arcsec<sup>-2</sup>, the individual  $q$  and PA radial profiles after their mean value has been subtracted. The solid thick line indicates the median of the variation at fixed radius, and the thick dashed lines enclose 68 per cent of all the curves. The profiles are remarkably flat, with little variation about the mean along a large radial range. We note that, because of this flatness, this behavior holds even if we



**Figure 3.** Radial variation about the mean of the isophotal axis ratio (top) and position angle (bottom) for low-luminosity Virgo galaxies. The light curves show the individual radial profiles, with their mean value subtracted. The solid thick line indicates the median of the variation at fixed radius, and the thick dashed lines enclose 68 per cent of all the curves. These low-luminosity systems exhibit almost flat axis ratio profiles, and no evidence for isophotal twisting.

scale the profiles by each individual effective radius. We however prefer this representation because it allows us to show the radial range that is potentially affected by the seeing PSF (shaded region in Fig. 3). The profiles show larger scatter at  $r < 2 \text{ FWHM}$ , with a slight preference for rounder shapes—consistent with mild circularization caused by seeing, as well as the presence of nuclear star clusters in some of these galaxies. In summary, low-luminosity quiescent Virgo galaxies are systems with remarkable isophotal simplicity and well described by a constant ellipticity parameter and major axis position angle. This lends support to the notion that homologous triaxial models are a good representation of their intrinsic shapes.

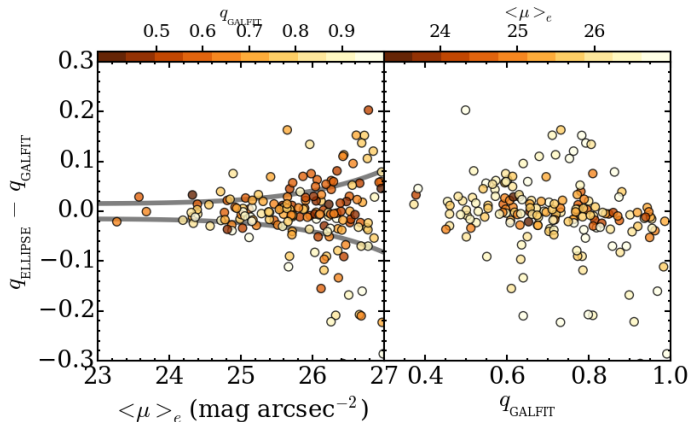
### 2.2.2. Accuracy of the axis ratio measurements



**Figure 4.** Error in the recovered axis ratios of simulated single-Sérsic galaxies as a function of their mean effective surface brightness (left panel), and axis ratio (right panel). The ordinates show the difference between the input axis ratios and the recovered GALFIT values, with the grey scale showing the density of galaxies in bins of  $\langle \mu \rangle_e$  and  $q$ . The long-dashed and short-dashed lines enclose 68 and 95 per cent of the data points, respectively.

Assessing the reliability of the axis ratio measurements for galaxies of such low central surface brightness is of paramount importance for this study, and we have carried out several consistency checks. First, as we have already mentioned, we simulated and added 36,500 single-Sérsic model galaxies to the images, sampling a range in structural parameters. The galaxy models are added to real NGVS frames after being convolved with the appropriate PSF for that randomly selected image location, and the simulated galaxy is fitted with GALFIT in the same way the real Virgo galaxies are. Figure 4 shows the difference between recovered and input axis ratios for the simulated galaxies as a function of their true mean effective surface brightness (left panel), and their true axis ratio (right panel). The grey scale shows the density of galaxies in bins of constant  $\langle \mu \rangle_{e,in}$  or  $q_{in}$ , with the solid line indicating the median offset. The long-dashed and short-dashed lines enclose 68 and 95 per cent of the data. It is clear that the axis ratios are very robustly determined down to mean effective surface brightnesses as low as  $\langle \mu \rangle_e \approx 28 \text{ mag arcsec}^{-2}$ , displaying no biases and a maximum  $1 \sigma$  error  $\approx 0.1$ . At fixed  $\langle \mu \rangle_e$  the error distribution is reasonably well described by a normal function. The right panel shows that the error distribution as a function of input axis ratio is highly peaked around zero. The bias toward recovering smaller axis ratios for larger  $q_{in}$  is only a natural consequence of the limitation that, by definition,  $q \leq 1$ . This is an often neglected but important bias, and in Sect. 3.2 we will discuss how to properly account for it in the analysis.

The uncertainties in  $q$  derived from these simulations can be regarded as lower limits only, in the sense that both the PSF and the galaxy light profile are perfectly described by the fitting functions, and the ellipticity is radially constant by construction. To obtain a better understanding of the amplitude of the measurement uncertainties, in Fig. 5 we compare the axis ratios from the ELLIPSE fits with those derived from the GALFIT models for the NGVS galaxies. The former measurements correspond to the mean  $q$  in the  $1 \leq r/r_e \leq 2$  range. The left and right panels in Fig. 5 show the axis ratio difference as a function of the GALFIT  $\langle \mu \rangle_e$  and  $q$ , respectively. The



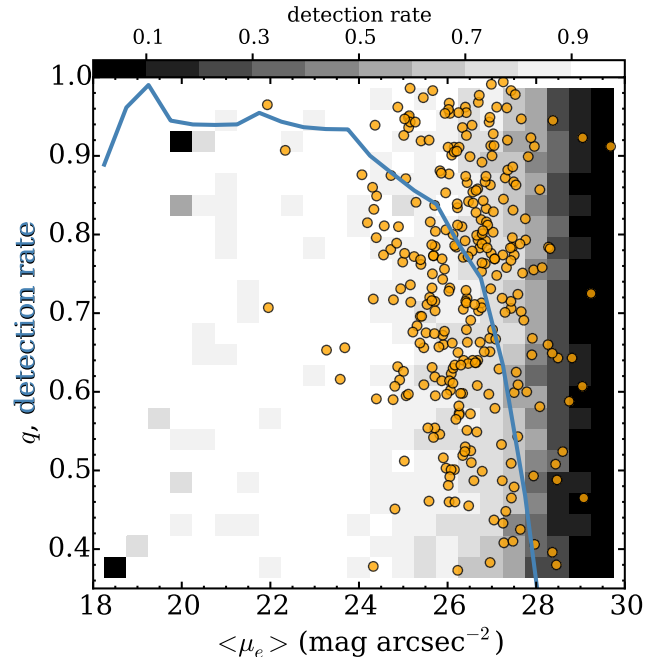
**Figure 5.** Difference between the mean axis ratios measured with ELLIPSE and the GALFIT  $q$  values for NGVS low-mass galaxies. The left panel shows differences as a function of their mean effective surface brightness, whereas the right panel shows them as a function of the GALFIT axis ratio measurements. Symbols are color-coded according to their  $q$  and  $\langle \mu \rangle_e$ , respectively. The solid grey lines in the left panel indicate the  $1\sigma$   $\langle \mu \rangle_e$ -dependent uncertainty we use in our analysis.

distributions of axis ratio differences are similar, both in shape and amplitude, to those in Fig. 4. The  $q$  values derived from these two completely different methods are remarkably consistent, and especially so for objects with high mean surface brightness. Of course, the disagreement increases at fainter  $\langle \mu \rangle_e$ , simply because this parameter is directly related to the mean signal-to-noise ratio across the galaxy body. It is also clear that there is no bias in the determination of axis ratios for either round or elongated galaxies.

In the analysis that follows we therefore use a  $\langle \mu \rangle_e$ -dependent uncertainty for the axis ratio measurements. At fixed  $\langle \mu \rangle_e$  the uncertainty follows a normal distribution, with a  $1\sigma$  dispersion that includes two components. The first one is the root-mean-square deviation of the ELLIPSE and GALFIT measurements,  $\sigma_q = 0.02$ , and represents the minimum uncertainty in the axis ratio estimates. This component is added quadratically to a  $\langle \mu \rangle_e$ -dependent  $\sigma_q$  computed through fits to the error distribution obtained from the simulations. The final  $1\sigma$  uncertainty is indicated by the solid lines in Fig. 5, and describes well the distribution of  $q$  differences. It is thus clear that the axis ratios for these faint Virgo galaxies are robustly determined irrespective of both the mean effective surface brightness and the apparent ellipticity of the galaxy.

### 2.3. Galaxy detection rate as a function of axis ratio and surface brightness

Under the assumption that these low mass galaxies are a family of optically thin triaxial ellipsoids (see Sect. 3.1), their observed surface brightness at fixed luminosity will depend on the intrinsic triaxiality and the orientation on the sky. For example, purely oblate (prolate) systems have brighter surface brightness when seen edge-on (face-on) as a result of larger integration of the emitted light along the line of sight. This can introduce a bias against small  $q$  values—if the population is oblate—or large  $q$  values—if prolate—in the distribution of observed axis ratios at surface brightnesses close to the detection limit of the survey.



**Figure 6.** Completeness as a function of mean effective surface brightness for simulated Sérsic galaxies. The grey scale shows the fraction of recovered simulated galaxies as a function of true mean effective surface brightness and true axis ratio. The solid line shows the detection rate as a function of  $\langle \mu \rangle_e$  only. NGVS reaches 50 per cent completeness at  $\langle \mu \rangle_e = 27.75$  mag arcsec $^{-2}$ . The symbols correspond to the distribution of real Virgo galaxies.

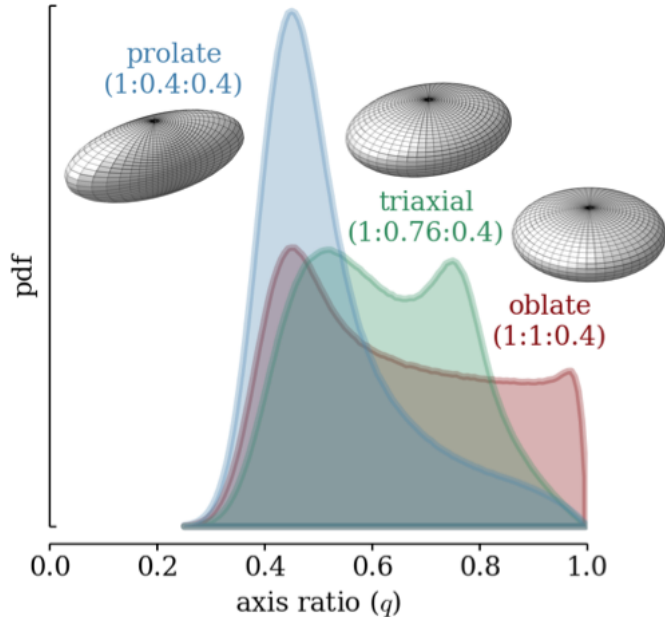
In Figure 6 we show, for the subsample of red-sequence low-mass galaxies, the distribution of measured axis ratios as a function of mean effective surface brightness. There is no obvious bias in the distribution of  $q$  as a function of  $\langle \mu \rangle_e$ . The background grey-scale image shows the detection rate for the sample of 36,500 simulated Sérsic galaxies as a function of these two parameters, and the solid line indicates the completeness as a function of  $\langle \mu \rangle_e$  only. The few bins showing high incompleteness at high surface brightness levels are caused by a combination of low number statistics, and geometrical non-detections—i.e., objects near the cores of very luminous galaxies, or galaxies that land in the vicinity of bright saturated stars. The NGVS reaches 50 per cent completeness at a remarkably faint  $\langle \mu \rangle_{e,lim} = 27.75$  mag arcsec $^{-2}$ . We prefer to be conservative, and in this work we further discard objects with  $\langle \mu \rangle_e$  values larger than this limit. We additionally discard five objects whose proximity to very bright stars or galaxies render their structural parameters highly uncertain. The final sample of Virgo low-mass galaxies in this study consists of 294 objects. We however note that the results of our analysis do not change if we include the fainter  $\langle \mu \rangle_e > \langle \mu \rangle_{e,lim}$  objects, which simply is a reflection of the unbiased distribution of  $q$  values, and the small subsample size.

## 3. BAYESIAN INFERENCE OF INTRINSIC SHAPES

### 3.1. A family of triaxial galaxy models

The ultimate goal of this study is to derive the underlying distribution of intrinsic shapes for faint galaxies in the Virgo cluster. While, strictly, full shape information for individual objects can only be derived by a combina-





**Figure 7.** Projected axis ratio distributions for three families of triaxial ellipsoids. The populations share a common mean intrinsic ellipticity  $\bar{E} = 0.6$ , as well as dispersions of ellipticities and triaxialities, namely  $\sigma_E = \sigma_T = 0.05$ . The families only differ in the mean triaxiality of the population  $\bar{T} = 0.01, 0.5$ , and  $0.99$  for the nearly oblate, triaxial, and nearly prolate populations, respectively. The mean intrinsic axis ratios  $1 : B/A : C/A$  for each family are indicated, together with a visualization of the corresponding three-dimensional ellipsoid.

tion of structural and kinematical data (e.g., Franx et al. 1991; Statler 1994; van den Bosch & van de Ven 2009), it is still possible to constrain the intrinsic shapes of a given population through the observed distribution of apparent axis ratios and a model for their three-dimensional shape. In the most general case, galaxies with similar elliptical isophotes are best described by triaxial models (Binney 1978). One can then try to invert the distribution of apparent axis ratios to infer the distribution of intrinsic shapes (Lambas et al. 1992; Lisker et al. 2007; Weijmans et al. 2014), but this deprojection problem has no unique solution unless some degree of axisymmetry is assumed—e.g., that the spheroids are either purely oblate, or purely prolate systems (Fall & Frenk 1983). Alternatively, one can try to reproduce the observed distribution of apparent axis ratios by varying the triaxial model parameters, and assuming the galaxies are randomly oriented on the sky (Hubble 1926; Sandage et al. 1970; Lambas et al. 1992; Ryden & Terndrup 1994; Sung et al. 1998; Vincent & Ryden 2005). This is the approach we follow in this work.

We model faint galaxies in Virgo as a family of optically thin triaxial ellipsoids. Following Franx et al. (1991), we assume that the three-dimensional galaxy density is structured as a set of coaligned ellipsoids characterized by a common ellipticity  $E = 1 - C/A$ , and a triaxiality  $T = (A^2 - B^2)/(A^2 - C^2)$ , where  $A \geq B \geq C$  are the intrinsic major, intermediate, and minor axis of the ellipsoid, respectively. Special cases of this model include the previously mentioned oblate ( $B = A$ , or  $T = 0$ ) and prolate ( $C = B$ , or  $T = 1$ ) spheroids. More general cases include the oblate-triaxial ( $C/B < B/A$ ) and prolate-

triaxial ( $C/B > B/A$ ) spheroids. In our modeling we further assume that the ellipticity and triaxiality parameters for our galaxy population are normally distributed (see Lambas et al. 1992; Ryden & Terndrup 1994; Ryden 2006; Holden et al. 2012), with means and standard deviations  $\bar{E}$ ,  $\sigma_E$ ,  $\bar{T}$ , and  $\sigma_T$ , respectively. Given values for the distributions of intrinsic axis ratios and random viewing angles for each model galaxy, one can easily compute the distribution of apparent axis ratios  $q$  (Stark 1977; Binney 1985).

Figure 7 shows three examples of apparent axis ratio distributions constructed from these ellipsoidal models for nearly oblate, nearly prolate, and triaxial families. In this example all models share a common mean intrinsic ellipticity  $\bar{E} = 0.6$ , as well as the same narrow dispersions of ellipticities and triaxialities, namely  $\sigma_E = \sigma_T = 0.05$ . The three models only differ in the mean triaxiality of the population, which takes values of  $\bar{T} = 0.01, 0.5$ , and  $0.99$  for the oblate, triaxial, and prolate families, respectively. Fig. 7 illustrates the dramatic impact that this difference makes in the distribution of projected axis ratios, especially at high  $q$ . While all three populations exhibit a peak around  $q = 1 - \epsilon \sim 0.4$  (associated with the intrinsic mean ellipticity  $\bar{E} = 1 - C/A = 0.6$ ), the relative number of apparently round objects progressively decreases as a function of  $\bar{T}$ . This is because the location of the secondary peak in the observed distribution of  $q$  is essentially determined by the mean  $B/A$ . A relative paucity of nearly round projected shapes is a generic feature of families of triaxial and prolate ellipsoids.

These apparent axis ratio distributions are equivalent to the likelihood of observing a given axis ratio given the model parameters, and we use this likelihood function to infer the intrinsic shape of low-luminosity galaxies using Bayesian statistics.

### 3.2. Bayesian analysis

Given some data  $D$  that we want to describe by a model with parameters  $\theta$ , for which we have some prior information  $I$ , Bayes' theorem states that:

$$p(\theta|D, I) = p(D|\theta, I) p(\theta|I) / p(D|I). \quad (1)$$

Here  $p(\theta|D, I)$  is the posterior probability density function (pdf) of the parameters  $\theta$  given the data  $D$  and prior information  $I$ ;  $p(D|\theta, I)$  is the likelihood function;  $p(\theta|I)$  is the prior joint probability of the model parameters in the absence of data; and  $p(D|I)$  is the marginal probability, i.e., the a priori probability of having the data marginalized over all possible parameters. In our particular case,  $D$  consists of the observed projected axis ratios for each individual galaxy,  $q_i$ ; the model parameters are  $\theta = \{\bar{E}, \sigma_E, \bar{T}, \sigma_T\}$ ; and the likelihood of apparent axis ratios is calculated as explained in Section 3.1. This approach has several formal and practical advantages. First, it does not require binning but works directly with discrete data, thus preserving the full information content of the dataset. This is always advantageous, but generally not mandatory when the inference is done on large samples of several thousand, or even hundreds of thousand objects (e.g., Ryden 2004; Holden et al. 2012). However, the analysis on discrete data becomes really powerful when the sample size is relatively small. Choos-

ing the optimal binning in these cases is not trivial, it always smears the underlying signal to a certain degree, and statistical uncertainties become non-Gaussian when the bins contain few objects.

In practice, in our Bayesian approach each distribution of axis ratios predicted by the models is convolved with another distribution representing the individual axis ratio  $q_i$  and its uncertainty. This is particularly useful when working with non-uniform and/or non-Gaussian uncertainties, or when the full probability distribution for the measurement of each individual axis ratio is available (e.g., Martin et al. 2008). For the very same reason, it is also straightforward to include the effects of biased measurements, and/or to account for lower or upper limits in the measured axis ratios—e.g. we only know our galaxy has  $q > 0.9$ . In this work we model the observed axis ratios as truncated normal distributions with means in the  $q \in [0, 1]$  interval, and dispersions that are dependent on each galaxy’s mean surface brightness, as discussed in Section 2.2.2. This naturally reproduces the bias towards lower recovered axis ratios at high true  $q$  also discussed in that Section. The excess probability from a regular normal distribution at  $q > 1$  (which is non-physical) gets redistributed towards  $q \leq 1$  values, creating a more extended tail that accounts for the fact that we tend to recover flatter shape measurements due to the bounded nature of axis ratios (see Fig. 4).

We use flat priors for the location parameters of our model normal distributions, i.e., the priors for  $\bar{E}$  and  $\bar{T}$  are uniformly distributed in the interval  $[0, 1]$ . For the standard deviations  $\sigma_E$  and  $\sigma_T$  we adopt scale-invariant priors of the form  $p(\sigma) \propto \sigma^{-1}$ . In order to efficiently sample the posterior distribution of our parameters  $\theta$  we use Markov Chain Monte Carlo (MCMC) methods. In particular, we use the affine-invariant MCMC algorithm implemented in the Python EMCEE package (Foreman-Mackey et al. 2013). We use a set of 100 ‘walkers’ that explore the parameter space at each of the 1000 steps we run. The MCMC chain converges to an equilibrium distribution after approximately 200 ‘burn-in’ steps, which we discard when exploring the posterior distributions.

#### 4. THE INTRINSIC SHAPES OF LOW-LUMINOSITY VIRGO GALAXIES

Figure 8 shows the MCMC posterior distributions for  $\bar{E}$ ,  $\sigma_E$ ,  $\bar{T}$ , and  $\sigma_T$  for our sample of faint satellites in Virgo. The panels on the diagonal show the marginalized posteriors for each parameter, and the dotted lines indicate the corresponding 2.5, 16, 50, 84 and 97.5 percentiles. The non-diagonal panels correspond to the joint posterior pdfs, and contours in these panels enclose the regions that contain 68 and 95 per cent of the cumulative posterior probability. There are several interesting results worth discussing in detail. First, the distribution of intrinsic ellipticities is well constrained by the data, with median and 68 per cent confidence values  $\bar{E} = 0.43_{-0.02}^{+0.02}$ , and  $\sigma_E = 0.11_{-0.01}^{+0.02}$ . The inferred triaxiality distribution of the population is  $\bar{T} = 0.16_{-0.06}^{+0.07}$  and  $\sigma_T = 0.04_{-0.03}^{+0.12}$ . It is clear that the triaxiality distribution is not well constrained with photometric data alone, something known since the early work from Binggeli (1980). Nevertheless, the data clearly favor nearly oblate models over nearly prolate ones, although in any case some degree of triax-

iality seems to be required. This is best appreciated in the marginalized posterior panel for  $\bar{T}$ : the maximum of the distribution is clearly different from zero, but fully triaxial and prolate shapes ( $\bar{T} \gtrsim 0.4$ ) are essentially ruled out by the data.

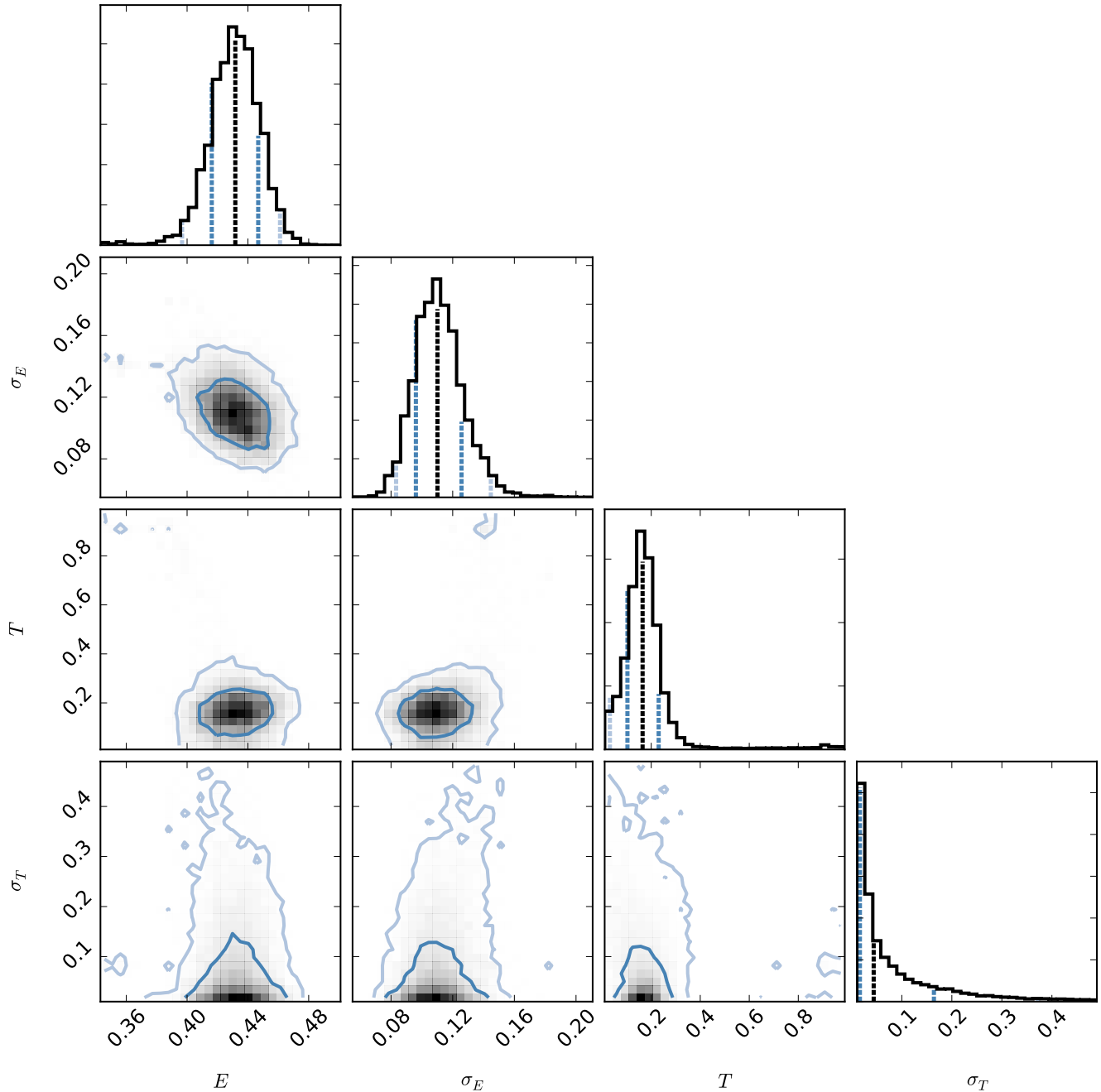
Figure 9 shows in grey a gaussian kernel density estimate of the observed axis ratio distribution. Each galaxy is represented by a truncated gaussian located at the corresponding measured  $q_i$ , and a with dispersion determined by its mean surface brightness (see Sect. 2.2.2). Thin lines correspond to 100 random samples from the MCMC chain, and the thick solid line indicates the model corresponding to the median of the samples in the marginalized distributions. For reference, the median estimates for  $\bar{E}$  and  $\bar{T}$  correspond to a triaxial ellipsoid with intrinsic axis ratios  $1 : 0.94 : 0.57$ —or, alternatively, it implies that faint Virgo galaxies are nearly oblate, thick spheroids.

The slight deviation from axisymmetry implied by  $\bar{T} \neq 0$  has to be taken with some caution. It is mostly driven by the secondary peak at  $q \sim 0.94$  in the observed distribution of axis ratios, and by the rapid decline in the number of even rounder galaxies. This relative paucity of round objects has been observed before, notably most recently by Weijmans et al. (2014) in their study of the intrinsic shapes of early-type galaxies in the ATLAS<sup>3D</sup> sample. The positive bias toward lower  $q$  introduced by the bounded nature of axis ratio measurements is too small of an effect to explain the lack of round objects in their sample of fast rotators. Perhaps even more puzzling is the fact that incorporating the kinematic misalignment information in their analysis resulted in a strong support for axisymmetry. Hence *real* triaxiality—understood as  $B/A \neq 1$  and a non-axisymmetric orbital structure—may not be the only explanation for the relative paucity of round galaxies. Alternative possibilities include the existence of a small  $m = 1$  mode, the presence of faint tidal features, or the suggestion that these galaxies may not be fully relaxed. We lack kinematic information for our NGVS low-mass galaxies, and therefore it is hard to say whether the small degree of triaxiality we infer is real, or related to the previous effect. In any case, if we impose pure oblateness as a prior ( $\bar{T} = \sigma_T = 0$ ) we infer an intrinsic flattening  $\bar{E} = 0.42 \pm 0.01$  and  $\sigma_E = 0.10 \pm 0.13$ , which are identical to the results from the full analysis—but in this case we of course cannot reproduce the observed  $q$  distribution at the roundest shapes.

We finally note that the models we are using implicitly assume that any obscuration by dust is independent of the viewing angle. This is a reasonable assumption, as we expect quiescent galaxies of these low masses to be optically thin systems with negligible amounts of dust (De Looze et al. 2010; Boselli et al. 2014). This is true for any low-mass, quiescent galaxy, but more so for these objects in the core of Virgo, where the stripping of the gas and dust components by ram-pressure from the intracluster medium occurs on short timescales (e.g., Boselli et al. 2008).

#### 5. COMPARISON WITH OTHER GALAXY SAMPLES AND SIMULATIONS

##### 5.1. Comparison with luminous early-type galaxies

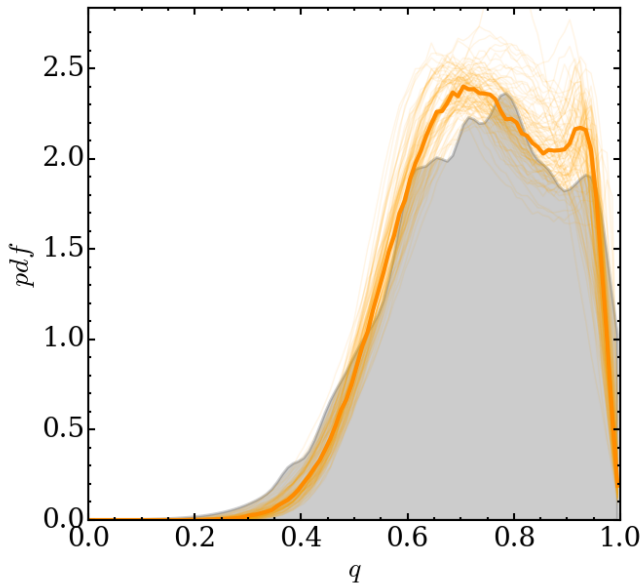


**Figure 8.** Posterior probability density functions for  $\bar{E}$ ,  $\sigma_E$ ,  $\bar{T}$ , and  $\sigma_T$  for faint satellites in Virgo. The panels in the diagonal show the posterior pdfs for each of the parameters, marginalized over all the other ones. The grey scale in the non-diagonal panels shows the corresponding joint posterior pdfs. Contours enclose the regions that contain 68 and 95 per cent of the cumulative posterior probability. The dotted lines in the diagonal panels indicate the 2.5, 16, 50, 84 and 97.5 percentiles of the corresponding marginalized posteriors.

The body of literature on the intrinsic shapes of early-type galaxies is enormous, and it is beyond the scope of this paper to present a comprehensive review. We will however discuss two important aspects of these studies: the shape dependence on stellar mass, and the constraints provided by kinematic studies.

In general, studies that make use of large photometric samples find that early-type galaxies have mean intrinsic minor-to-major axis ratios in the 2:3 to 2:5 range. Differences arise because of selection effects, including the definition of what is exactly an early-type galaxy, as well as

the mass range under study (Lambas et al. 1992; Vincent & Ryden 2005; Padilla & Strauss 2008). Holden et al. (2012) find that at the high mass end ( $M_\star > 10^{11} M_\odot$ ) intrinsically flattened quiescent galaxies are almost absent—they only constitute  $\sim 10$  per cent of the population—and that these systems are best described as a family of triaxial spheroids with  $\bar{T} = 0.45$  and  $\bar{E} = 0.37$ . This leads them to suggest that  $M_\star > 10^{11} M_\odot$  sets a mass scale above which disk-destroying multiple major mergers are the dominant method of mass assembly (van der Wel et al. 2009). The fraction of oblate spheroids in-



**Figure 9.** The grey area shows a gaussian kernel density estimate of the observed axis ratios of faint galaxies in Virgo. The model corresponding to the median of the samples in the marginalized distributions is shown as a thick solid line, and corresponds to a family of nearly oblate spheroids with mean intrinsic ellipticity  $\bar{E} = 0.43$ . Thin lines correspond to 100 random samples from the MCMC chain.

creases toward lower masses, where disks or flattened ellipticals populate the mass function. Note that pure prolate shapes have long been ruled out for this population of massive early-type galaxies (e.g., Saglia et al. 1993).

This result is perhaps best understood by separating early-types into distinct kinematical classes. As discussed above, Weijmans et al. (2014) combined photometric and kinematic data for early-type galaxies in the ATLAS<sup>3D</sup> survey to carry out a comprehensive investigation of their intrinsic shapes. They find that fast rotators (FRs) are much more flattened than slow rotators (SRs). The former population is predominantly axisymmetric, with intrinsic shapes similar to those of spiral galaxies, whereas the latter are best described as mildly triaxial systems. The relative abundance of these two families as a function of mass—with SRs dominating at higher masses—naturally explains the prevalence of rounder shapes at the highest masses found by the photometric studies mentioned above.

### 5.2. Comparison with luminous dEs in Virgo

The first statistical study on the intrinsic flattening of Virgo quiescent galaxies was carried out by Ichikawa (1989) using a sample of nearly one hundred galaxies in the  $-17 < M_B < -13$  magnitude range. This early study already showed that Virgo dEs are best described as a population of nearly oblate spheroids with mean intrinsic thickness  $\approx 0.6$ . Ryden & Terndrup (1994) modeled the intrinsic distribution of axis ratios for a sample of 70 Virgo dEs in the  $-17 \lesssim M_B < -14$  mag range. They found mean intermediate-to-major and minor-to-major axis ratios of  $B/A = 0.95$  and  $C/A = 0.54$ , re-

spectively. These values are remarkably similar to ours for the fainter population; dEs in their sample also are thick, nearly oblate spheroids. We however note that the subclass of bright nucleated dEs was found to be significantly rounder ( $C/A = 0.88$ ) than the non-nucleated dEs ( $C/A = 0.45$ ; see also Ferguson & Sandage 1989). Here we do not attempt such separation, but it will be explored in a future paper in this series where we extend the analysis to the entire Virgo cluster—hence improving the statistics.

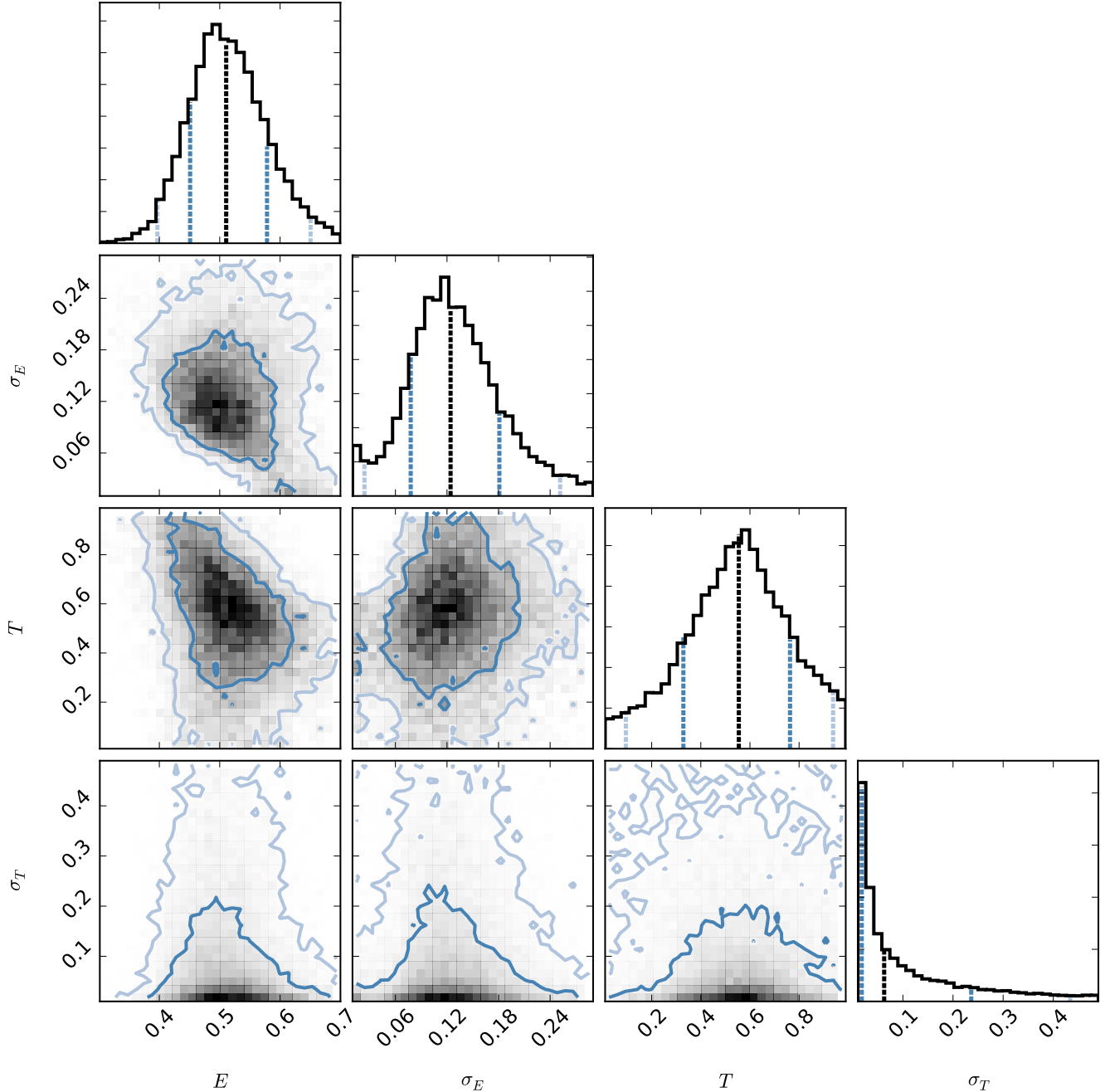
More recently, Lisker et al. (2007) studied a much larger sample of  $\sim 400$  Virgo dEs as faint as  $M_B = -13$  mag using SDSS imaging data. They inferred the intrinsic axis ratio distribution by inverting the observed  $q$  distributions under the assumption that the population is either purely oblate, or purely prolate. They also distinguished between nucleated and non-nucleated systems and, like Ryden & Terndrup (1994) before, found that the latter are significantly flatter than the former. In both cases, though, oblate shapes are favored over prolate ones. For the class of nucleated dEs they found a wide distribution of intrinsic thickness, ranging from 0.5 to 0.8 irrespective of luminosity. They found even more flattened shapes for dEs that show evidence for the presence of spiral arms or residual star formation, suggesting that these systems are probably disk-like in nature (see also Aguerri et al. 2005).

As discussed above, kinematic data, and especially integral-field observations, are crucial to advance intrinsic shape studies. There is a growing amount of such data for luminous cluster dEs (e.g., Toloba et al. 2014; Ryś et al. 2014), and it would be desirable to augment even more the sample size so that kinematics can be incorporated into the inference process. This will not only help on constraining axisymmetry, but also in separating the dEs into distinct kinematical classes. For very faint systems like the ones studied here, however, we will have to rely on building statistically large photometric samples to constrain their shapes.

### 5.3. Comparison with Local Group satellites

The Local Group provides a natural comparison sample, because its satellite system comprises objects of similar intrinsic physical properties—e.g., luminosity, gas content, or current star formation rate—to those in Virgo. While low number statistics made the task of deriving the intrinsic shapes of LG satellites almost insurmountable in the past, the advent of large panoramic surveys in the last decade has boosted the number of known satellites around both the MW (e.g., Koposov et al. 2008), and M31 (e.g., Richardson et al. 2011).

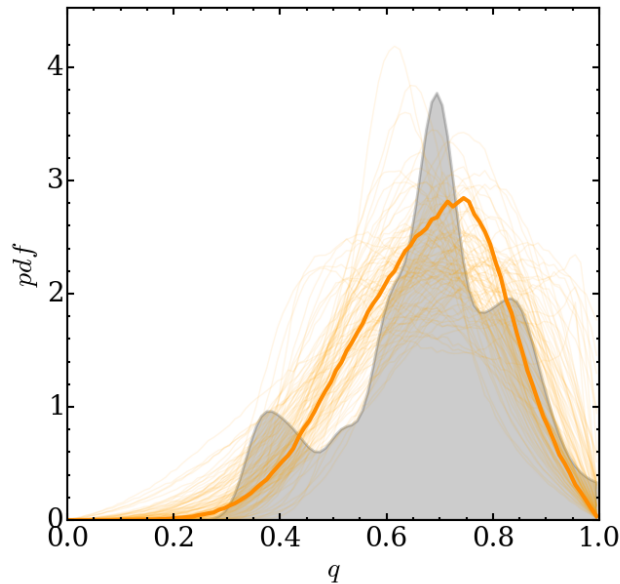
We carry out an investigation on the intrinsic shapes of LG satellites using the machinery developed in Sect. 3. We use the structural information for low-luminosity galaxies in the nearby Universe compiled by McConnachie (2012). In order to select a sample of LG satellites comparable to the ones in Virgo, we introduce a series of cuts. First, we limit the sample to actual LG satellites by selecting only low-luminosity galaxies within the zero-velocity surface of the LG (see fig. 5 and discussion in McConnachie 2012). Second, we select ‘classical’ satellites having  $-14 \leq M_V \leq -8$  mag, thus matching our luminosity range for Virgo systems. Note that restricting the analysis to that luminosity in-



**Figure 10.** Same as Fig. 8, but for Local Group dSphs. The panels show the posterior probability density functions for  $\bar{E}$ ,  $\sigma_E$ ,  $\bar{T}$ , and  $\sigma_T$  for faint satellites in the LG. The panels in the diagonal show the posterior pdfs for each of the parameters, marginalized over all the other ones. The grey scale in the non-diagonal panels shows the corresponding joint posterior pdfs. Contours enclose the regions that contain 68 and 95 per cent of the cumulative posterior probability, respectively. The dotted lines on the diagonal panels indicate the 2.5, 16, 50, 84 and 97.5 percentiles of the corresponding marginalized posteriors.

terval is further justified given the claimed discontinuity in apparent dSph shapes below  $M_V \approx -7.5$  (Martin et al. 2008). Finally, we discard all systems with  $M_{\text{HI}}/L_V > 0.1 M_\odot/L_\odot$ . By limiting our sample to objects with low gas mass fractions—or, equivalently, little current star formation activity (e.g., Weisz et al. 2011)—we better mimic our selection of red-sequence faint cluster galaxies. Satellites with a significant neutral gas component can have different intrinsic shapes than their gas-poor counterparts (e.g., Roychowdhury et al. 2013),

which is natural in a scenario where young stars form within the highest density regions of a kinematically cold HI disk. With this cut we avoid potential biases in our analysis, and it results in a final sample of 23 LG satellites. By combining galaxies from both M31 and the MW, we mitigate to some degree the potential complications arising from observing the latter population from a preferred viewing angle (Barber et al. 2015). The axis ratio measurements for most M31 satellites come from the homogeneous statistical analysis by Martin et al. (2015).



**Figure 11.** Same as Fig. 9, but for Local Group dSphs. The grey area shows a kernel density estimate of the observed axis ratios of faint satellites in the LG. The model corresponding to the median of the samples in the marginalized distributions is shown as a thick solid line, and corresponds to a family of triaxial spheroids with mean intrinsic ellipticity  $\bar{E} = 0.51$ . Thin lines correspond to 100 random samples from the MCMC chain.

For the remaining galaxies, we use the  $q$  values and uncertainties as compiled by McConnachie (2012).

Figure 10 shows the MCMC posterior distributions for  $\bar{E}$ ,  $\sigma_E$ ,  $\bar{T}$ , and  $\sigma_T$  for the final sample of 23 LG satellites. It is immediately clear that the much smaller sample size translates into poorer constraints on their intrinsic shapes compared to the Virgo sample. Nevertheless, the intrinsic ellipticity is relatively well constrained, and the mean and dispersion of the population distribution are remarkably similar to those of faint Virgo galaxies:  $\bar{E} = 0.51^{+0.07}_{-0.06}$  and  $\sigma_E = 0.12^{+0.06}_{-0.05}$ , respectively. The intrinsic triaxiality of the LG population, however, is significantly different from that of the Virgo sample. LG satellites are best reproduced by triaxial models ( $\bar{T} = 0.55^{+0.21}_{-0.22}$ ).

The difference in mean triaxiality between the LG and Virgo low-mass galaxies can be readily understood with the help of Fig. 11. As before, the grey area shows a kde estimate of the observed axis ratios in the LG. Thin lines again correspond to 100 random samples from the MCMC chain, and the thick solid line indicates the model corresponding to the median of the samples in the marginalized distributions. The relevant feature of this figure is the marked paucity of objects with very round apparent shapes. As discussed in Sect. 3.1 and illustrated in Fig. 7, this is a characteristic of intrinsically triaxial or prolate populations, simply because random projections on the sky of these objects rarely result in perfectly circular shapes.

The inferred mean values for the model distributions correspond to a triaxial ellipsoid with intrinsic axis ratios 1 : 0.76 : 0.49. We note that purely oblate shapes ( $\bar{T} = 0$ ) provide a poor description of the observed axis ratio

distribution. Purely prolate shapes ( $\bar{T} = 1$ ), on the other hand, do a better job at reproducing the distribution in Fig. 11, but result in thicker spheroids,  $\bar{E} = 0.41 \pm 0.04$  and  $\sigma_E = 0.13^{+0.04}_{-0.03}$ . This is already evident in the  $\bar{E}$ - $\bar{T}$  panel of Fig. 10, where the contours enclosing 95 per cent of the posterior probability extend toward less flattened shapes at high  $\bar{T}$ .

Martin et al. (2008) presented a comprehensive analysis of the structural properties of faint satellites around the MW, including their *apparent* ellipticities. They found that objects in the  $-14 \lesssim M_V \lesssim -7.5$  magnitude range (essentially the very same ones used here) have relatively round shapes, with a mean observed  $q \approx 0.7$  (compare with the peak of the observed distribution in Fig. 11). Remarkably, there seems to be a transition in galaxy shape around this luminosity, as fainter systems have more elongated shapes (mean  $q \sim 0.5$ ), and a broader distribution of apparent axis ratios. They explore several possibilities for the origin of these very elongated shapes, including flattening by rotation, a more elongated shape of their dark matter halos, and tidal distortions. They argue that the latter scenario—that tidal interactions with the MW drive dSph shapes—is slightly preferred over the other mechanisms, and suggest that the broad distribution of observed ellipticities ( $0.2 \lesssim \epsilon \lesssim 0.8$ ) may correspond to different stages in the disruption process (cf. Sect. 5.4).

Recently, Salomon et al. (2015) have inferred the intrinsic ellipticity of M31 satellites under the assumption that their visible components have prolate shapes, and after exploring different possibilities for the orientation of the galactic major axis. They find that the population has a mean intrinsic axis ratio  $C/A \sim 0.5$ , which is virtually identical to the value derived here. A non-negligible fraction of the studied satellites (6/25) are found to be intrinsically significantly elongated, and these authors suggest they may be strongly affected by tides. This, in turn, would have important implications for the orbital distribution of M31 dSphs, because significant tidal stretching is known to be a short-lived phase mostly limited to pericentric passages (Barber et al. 2015).

It would be extremely interesting to test the disruption hypothesis in the case of extremely faint Virgo galaxies because they reside in a much more dense environment than the MW, and have done so for longer times. Recently, Jang & Lee (2014) reported the serendipitous discovery of a candidate ultra-faint ( $M_V \approx -6.5$ ) galaxy in very deep HST/ACS images of the Virgo cluster. The fact that the apparent ellipticity of this one object is much lower ( $\epsilon \approx 0.1$ ) than what is generally found in the LG opens up the exciting possibility that these ultra-faint systems have different intrinsic shapes in groups and clusters. This is admittedly pure speculation, as one single object hardly constrains the intrinsic shape of a population. But, unfortunately, relatively large samples of such galaxies in this very low-mass regime will remain inaccessible in Virgo for years to come.

#### 5.4. Comparison with numerical simulations of tidal evolution for satellites

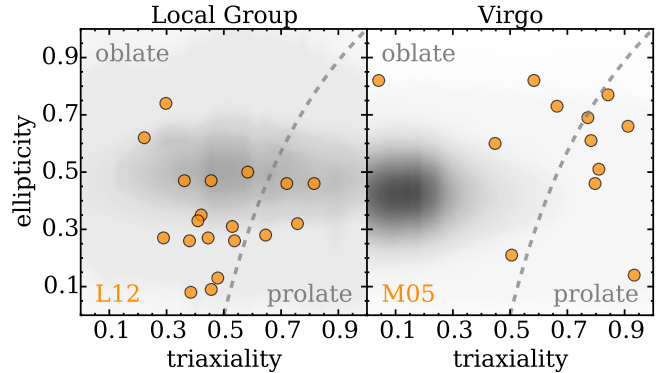
Tidal interactions between a satellite and its host, as well as with other subhalos, are essentially inevitable during their orbital evolution. Numerical simulations show

that tidal stripping and shocking generally result in significant mass loss from the dark matter halo (e.g. Gnedin 2003; Peñarrubia et al. 2008; Smith et al. 2013b). Depending on the strength and duration of the interaction, and on the properties of the progenitor galaxy, tidal effects can also drive significant morphological transformation, redistribution of the stellar orbits, and even inflict stellar mass losses.

Barber et al. (2015) studied the shapes of the gravitational potential of subhalos in the Aquarius simulations at radii comparable to the typical half-light radii of LG dSphs. They show that inner isopotentials are nearly prolate, with intrinsic axis ratios  $B/A \approx 0.8$  and  $C/A \approx 0.75$ . While departure from sphericity is in good agreement with our results for the LG satellites, the mean thickness inferred from observations,  $C/A \approx 0.5$ , is significantly lower than that of simulated subhalos (see also Salomon et al. 2015). The cumulative effect of continuous tidal stripping, however, is to reduce the asphericity of the subhalos. It is tempting to associate this effect to the rounder shapes we have inferred for Virgo low-luminosity galaxies compared to LG dSphs. Not only is tidal stripping stronger in the cluster potential, but a large fraction of cluster galaxies have spent many Gyr orbiting within group-sized subhalos before being accreted onto the cluster (Lisker et al. 2013). For this to be the case, however, requires the host subhalo of these galaxies to be heavily stripped ( $> 85 - 90\%$  of the initial dark matter mass; see Smith et al. 2013b, 2015; Mistani et al. 2016). Strong tidal stripping can occur as the satellite galaxies orbit a more massive host during their infall onto the cluster (e.g., Paudel et al. 2013), or by the cluster potential itself. In fact, the field under study contains some obvious cases of tidally stripped low-luminosity galaxies—like VCC1297, which is a prototype for the more luminous cE class (see also Guérou et al. 2015).

An even more dramatic transformation, in terms of both mass loss and redistribution of stellar orbits, can occur if the progenitor galaxy is a dynamically cold disk. In this case, the cumulative effects of tidal stripping, tidal shocking, and resonant stripping result in the creation of a lower mass, spheroidal-like system (Moore et al. 1996, 1998; Mayer et al. 2001; Mastropietro et al. 2005; D’Onghia et al. 2009; Smith et al. 2010; Kazantzidis et al. 2011). Naturally, these mechanisms have been invoked to account for the morphology-density relation in both galaxy-sized and cluster-sized halos.

In Figure 12 we compare the intrinsic shapes of faint galaxies in the LG and Virgo inferred in the present work with the predictions from two sets of numerical simulations of tidally transformed disks. For the comparison with the LG satellites we use the set of simulations from Lokas et al. (2012), who follow 10 Gyr of orbital evolution of a galaxy with initial stellar mass  $M_\star \approx 2 \times 10^7 M_\odot$  within an MW-sized halo. The final intrinsic ellipticity and triaxiality of the remnants is shown with round symbols in Fig. 12 (left panel). The shaded region corresponds to 1000 random bivariate normals from the MCMC chain for the model parameters of LG satellites (Sect. 5.3). The dashed line marks the transition between prolate-triaxial and oblate-triaxial spheroids. There is reasonably good agreement between the data and the simulations, with the majority of the simulated remnants



**Figure 12.** Comparison between the inferred intrinsic ellipticity and triaxiality for faint galaxies in the Virgo cluster and in the Local Group, and the predictions from numerical simulations of tidally modified disks orbiting cluster- and galaxy-sized halos, respectively. In both panels, the shaded region corresponds to 1000 random bivariate normals from the MCMC chain for the model parameters. The dashed line marks the transition between prolate-triaxial and oblate-triaxial spheroids. In the left panel, the symbols show the results for the set of simulations of low-mass disks evolved in a MW-like potential by Lokas et al. (2012). In the right panel we show the final shapes of discs evolved within a Virgo-like halo (Mastropietro et al. 2005). Note that, while the LG simulations roughly match the inferred distribution of intrinsic shapes, the remnants simulated in the cluster environment are both too flattened and too prolate compared to observed faint Virgo galaxies (but see text for discussion).

being oblate-triaxial spheroids with moderate to small intrinsic ellipticities. The three simulated galaxies with the roundest shapes—which are the most tidally stripped systems—are the least consistent with the distribution of intrinsic shapes inferred in this work. It is however important to note that, in their set of simulations, Lokas et al. (2012) vary not only the orbital parameters of the satellite, but also their overall stellar and dark matter halo properties (including the initial stellar thickness and mass, and the dark halo concentration, spin, and mass). Their simulations can only be interpreted as being representative of the range of shapes one can expect from the evolution under the tidal stirring scenario, and never as predicting a full distribution of realistic shapes. As discussed before, the paucity of observed LG satellites with very round apparent shapes strongly suggests that the population must have a certain degree of intrinsic triaxiality.

We compare the intrinsic shapes of low-luminosity Virgo galaxies with the set of simulations from Mastropietro et al. (2005). These authors follow the evolution of a Large Magellanic Cloud-type disk galaxy infalling into a Virgo-sized halo at  $z = 0.5$ . Their simulated galaxies undergo significant mass loss and morphological evolution, and their final intrinsic shapes are shown in the right panel of Fig. 12. It is clear that the remnants from this set of simulations do not resemble the population of faint galaxies in the core of Virgo: they are both too flattened and too prolate compared to the inferred intrinsic shapes of low-luminosity galaxies in the NGVS. We note that in Fig. 12 we only include the simulated remnants that by  $z = 0$  have experienced at least one orbit with a pericentre bringing them within the virial radius.

However, this striking difference is probably just a limitation of using this particular set of simulations for the

comparison. The galaxies simulated by M05 are initially too massive to be the progenitors of the faint Virgo galaxies in our sample: even their most heavily stripped remnant has a final stellar mass at  $z = 0$  roughly seven times larger than our most massive object. This relatively high initial mass ( $M_* \approx 4 \times 10^9 M_\odot$ ) also implies a high disk self gravity, which coupled to a very thin and cold initial configuration makes the galaxy rather bar-unstable. Indeed, the severe morphological transformation experienced by the galaxy in these simulations is essentially driven by the formation of a strong bar in the early stages of its tidal evolution. The low surface density material remaining around the bar is subsequently easily stripped by tidal forces, leaving behind a prolate-triaxial stellar body (essentially, a naked bar). The bar usually buckles, and the system becomes progressively rounder as additional tidal interactions and heating occur.

It is important to note that the model galaxy in M05 is so bar-unstable that *all* of their twenty simulated galaxies form a bar at some point during their evolution—including all of the less-perturbed galaxies that do not even penetrate the cluster virial radius during their orbital evolution (see their table 1, and the discussion in section 3 of that paper). This is probably unrealistic. While it is true that the bar fraction *in the field* peaks at about 50 per cent around a stellar mass  $M_* \sim 10^9 M_\odot$ , only one out of five Virgo disk galaxies are barred in this stellar mass regime—and the bar fraction quickly drops to zero at even lower masses (Méndez-Abreu et al. 2010, 2012). This suggests that bar-driven evolution due to tidal interactions in the cluster environment is less important than these simulations predict. This scenario is fully consistent with the lack of prolate, bar-like satellites in Virgo.

It is possible that the apparent discrepancy between the inferred intrinsic shapes of faint Virgo galaxies and those predicted by numerical simulations of tidal interactions can be reconciled simply by simulating more adequate models for their progenitor galaxies. For instance, it is clear now that ‘central’ (as opposed to satellite) galaxies with masses  $M_* < 10^9 M_\odot$  are not thin disks, but rather thick oblate spheroids (Sánchez-Janssen et al. 2010; Roychowdhury et al. 2013) in which pressure support is dominant (Wheeler et al. 2015). Additionally, Kazantzidis et al. (2011) show that tidal shocks heat the stellar body more effectively in lower mass systems, and as a result the remnants evolve into more spherical shapes. Note, however, that these simulations do not succeed in fully removing rotation, but rather decrease  $v/\sigma$  by increasing the stellar velocity dispersion. Therefore, they do not provide a perfect description of the faintest LG satellites—which lack rotation support—, but transformation through tidal stirring may be an appropriate scenario for more massive dE-type galaxies. In any case, and to our best knowledge, no simulations have yet studied the tidal evolution of galaxies in the low luminosity regime explored in this work within a Virgo-like environment.

##### 5.5. Comparison with low-luminosity, star-forming galaxies in the field

Finally, our results can also be compared with the intrinsic shape of faint star-forming galaxies, i.e., similar mass objects residing in lower density environments

and having later-type morphologies and measurable star formation rates. Sung et al. (1998) already noted that faint star-forming galaxies do not conform to the belief that they are highly flattened systems; much on the contrary, they actually are relatively puffy objects. Sánchez-Janssen et al. (2010) later showed that, under the assumption that low-mass galaxies are well described as oblate spheroids, their thickness systematically increases towards fainter luminosities. They argued this trend to be an intrinsic property—faint galaxies in their sample are mostly centrals, not satellites—related to the complex interplay between mass, specific angular momentum, and stellar feedback effects. Simply put, the increasing importance of turbulent motions over rotational support in lower mass galaxies leads to the formation of thicker systems (e.g. Kaufmann et al. 2007). The recent kinematical analysis of low-mass galaxies in the Local Volume by Wheeler et al. (2015) provides strong support for this scenario. These authors find that the great majority of  $M_* < 10^8 M_\odot$  galaxies, regardless of whether they are isolated or satellites, are pressure-supported.

A more quantitative comparison with the results presented here can be made by using the recent analysis of Roychowdhury et al. (2013). They follow a similar approach to ours to derive the intrinsic shapes of a sample of several hundred star-forming faint galaxies in the Local Volume. Their optical axial ratios were measured at the Holmberg isophote,  $\mu_B = 26.5 \text{ mag arcsec}^{-2}$ , and consequently they should be a fair representation of the underlying galaxy shape rather than being affected by recent star formation events. These authors find the same trend as Sánchez-Janssen et al. (2010), namely that fainter star-forming systems are systematically rounder. Their lowest-luminosity subsample ( $-13 \lesssim M_B \lesssim -7 \text{ mag}$ ) is a reasonably good match to the luminosities of the faint Virgo systems studied in this work. These nearby dIrrs are best described by a family of oblate spheroids with mean intrinsic ellipticity  $\bar{E} = 0.52$ . This is again a very similar value—though implying slightly flatter shapes—to the one we derive here for both Virgo and, especially, LG quiescent galaxies of the same luminosity. This finding strongly indicates that, to first order, the mechanisms shaping the structure of these low-mass systems are almost independent of the *present* local environment. If the progenitors of low-luminosity Virgo galaxies did have shapes similar to present-day faint dIrrs—an assumption that does not necessarily have to hold—then the environmental effects that transformed them into non-star forming galaxies have resulted in a relatively small amount of thickening of their stellar bodies.

## 6. SUMMARY AND CONCLUSIONS

We have investigated the intrinsic shapes of faint, red-sequence galaxies in the central 300 kpc of the Virgo cluster using deep imaging obtained as part of the NGVS. We have built a sample of nearly 300 red-sequence cluster members in the  $-14 < M_g < -8$  magnitude range, and we measure their apparent axis ratios through Sérsic fits to their two-dimensional light distribution in the  $g$ -band. The resulting distribution of apparent axis ratios is then fit by a family of triaxial models with normally-distributed intrinsic ellipticities and triaxialities. We develop a Bayesian framework to explore the posterior distribution of the model parameters. This approach al-



lows us to work directly on discrete data (i.e., without binning), as well as to account for individual, surface brightness-dependent axis ratio uncertainties.

We find that faint galaxies in the core of Virgo are best described as thick, nearly oblate spheroids, with mean intrinsic axis ratios  $1 : 0.94 : 0.57$ . Using the machinery developed for this investigation, we additionally carry out a study of the intrinsic shapes of satellites of similar luminosities in the Local Group. For this population we infer a slightly larger mean intrinsic ellipticity, but the lack of objects with round appearances translates into more triaxial shapes,  $1 : 0.76 : 0.49$ . We compare these results with numerical simulations that follow the tidal evolution of satellites within cluster- and group-sized halos. The LG simulations generally match the inferred distribution of dSph intrinsic shapes, but the remnants simulated in the cluster environment are both too flattened and too prolate compared to observed faint Virgo galaxies. We discuss possible reasons for this discrepancy, and conclude that more adequate simulations of cluster galaxies in the luminosity regime explored in this work are needed in order to compare with the observations.

We finally compare the intrinsic shapes of the NGVS low-luminosity galaxies with samples of more massive, quiescent cluster systems, and with star-forming galaxies of similar luminosities residing in lower-density environments. Remarkably, all these low-mass galaxy samples are also best described as thick, nearly oblate spheroids, and their intrinsic ellipticities and triaxialities are consistent, within the uncertainties, with those inferred for the faint Virgo population. This is an important result, as it implies that, to first order, *intrinsic shapes in this luminosity regime are largely independent of the environment in which the galaxy resides.*

There is, however, a hint that cluster objects may be slightly thicker than group satellites, which are not different from more isolated faint galaxies. Indeed, the mean thickness of the low-luminosity population is  $C/A = 0.57, 0.49, \text{ and } 0.48$  for objects in Virgo, in the LG, and in the field, respectively. The fact that low-mass galaxies that have a priori experienced very different degrees of environmental effects have very similar flattenings seems to suggest that internal—and not external—mechanisms are the main drivers of galaxy structure in this luminosity regime. This may be the case if, for example, galaxies of these masses already form as thick, puffy systems due to baryonic processes that act to pressurize gas within low-mass halos. These processes may include photoionization from the cosmic UV background, and star formation feedback (Kaufmann et al. 2007; Teyssier et al. 2013). Alternatively, other sources of internal heating, such as dynamical interactions with dark subhalos may play a role as well (Starkenburg & Helmi 2015). Subsequent gas stripping and tidal evolution within their host halos—especially if acting early on—certainly can contribute to some degree to the heating of the stellar body of satellites, making them progressively thicker. However, we cannot rule out the possibility that these systems were already formed with the observed distributions of intrinsic flattenings.

We will explore this scenario in more detail in a future paper of this series using an extended sample of several thousand faint NGVS galaxies across the entire Virgo cluster, for which we are currently deriving structural

parameters and establishing robust membership. With increased statistics we will be able to investigate intrinsic shapes as a function of mass, color, and location in position-velocity phase-space, thus addressing the interplay between internal and external processes in shaping the structure of low-luminosity galaxies.

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## REFERENCES

- Aguerri, J. A. L., Iglesias-Páramo, J., Vílchez, J. M., Muñoz-Tuñón, C., & Sánchez-Janssen, R. 2005, *AJ*, 130, 475  
 Barber, C., Starkenburg, E., Navarro, J. F., & McConnachie, A. W. 2015, *MNRAS*, 447, 1112  
 Benson, A. J., Bower, R. G., Frenk, C. S., et al. 2003, *ApJ*, 599, 38  
 Bertin, E., & Arnouts, S. 1996, *A&AS*, 117, 393  
 Bialas, D., Lisker, T., Olczak, C., Spurzem, R., & Kotulla, R. 2015, *A&A*, 576, A103  
 Binggeli, B. 1980, *A&A*, 82, 289  
 Binney, J. 1978, *MNRAS*, 183, 501  
 —. 1985, *MNRAS*, 212, 767  
 Blakeslee, J. P., Jordán, A., Mei, S., et al. 2009, *ApJ*, 694, 556

- Boselli, A., Boissier, S., Cortese, L., & Gavazzi, G. 2008, *ApJ*, 674, 742
- Boselli, A., Cortese, L., Boquien, M., et al. 2014, *A&A*, 564, A67
- Boselli, A., & Gavazzi, G. 2014, *A&A Rev.*, 22, 74
- Brooks, A. M., Kuhlen, M., Zolotov, A., & Hooper, D. 2013, *ApJ*, 765, 22
- Bullock, J. S., Kravtsov, A. V., & Weinberg, D. H. 2000, *ApJ*, 539, 517
- Chiboucas, K., Tully, R. B., Marzke, R. O., et al. 2011, *ApJ*, 737, 86
- Côté, P., Piatek, S., Ferrarese, L., et al. 2006, *ApJS*, 165, 57
- De Looze, I., Baes, M., Zibetti, S., et al. 2010, *A&A*, 518, L54
- Dekel, A., & Silk, J. 1986, *ApJ*, 303, 39
- den Brok, M., Peletier, R. F., Seth, A., et al. 2014, *MNRAS*, 445, 2385
- D’Onghia, E., Besla, G., Cox, T. J., & Hernquist, L. 2009, *Nature*, 460, 605
- Drinkwater, M. J., Jones, J. B., Gregg, M. D., & Phillipps, S. 2000, *Publications of the Astronomical Society of Australia*, 17, 227
- Duc, P.-A., Cuillandre, J.-C., Karabal, E., et al. 2015, *MNRAS*, 446, 120
- Durrell, P. R., Côté, P., Peng, E. W., et al. 2014, *ApJ*, 794, 103
- El-Badry, K., Wetzell, A. R., Geha, M., et al. 2015, *ArXiv e-prints*, arXiv:1512.01235
- Fall, S. M., & Frenk, C. S. 1983, *AJ*, 88, 1626
- Ferguson, H. C., & Sandage, A. 1989, *ApJ*, 346, L53
- Ferrarese, L., Côté, P., Cuillandre, J.-C., et al. 2012, *ApJS*, 200, 4
- Foreman-Mackey, D., Hogg, D. W., Lang, D., & Goodman, J. 2013, *PASP*, 125, 306
- Franx, M., Illingworth, G., & de Zeeuw, T. 1991, *ApJ*, 383, 112
- Geha, M., Blanton, M. R., Yan, R., & Tinker, J. L. 2012, *ApJ*, 757, 85
- Gnedin, O. Y. 2003, *ApJ*, 582, 141
- Governato, F., Brook, C., Mayer, L., & et al. 2010, *Nature*, 463, 203
- Grebel, E. K., Gallagher, III, J. S., & Harbeck, D. 2003, *AJ*, 125, 1926
- Guérou, A., Emsellem, E., McDerimid, R. M., et al. 2015, *ApJ*, 804, 70
- Haşegan, M., Jordán, A., Côté, P., et al. 2005, *ApJ*, 627, 203
- Hargis, J. R., Willman, B., & Peter, A. H. G. 2014, *ApJ*, 795, L13
- Helmi, A., Sales, L. V., Starkenburg, E., et al. 2012, *ApJ*, 758, L5
- Hilker, M., Infante, L., Vieira, G., Kissler-Patig, M., & Richtler, T. 1999, *A&AS*, 134, 75
- Holden, B. P., van der Wel, A., Rix, H.-W., & Franx, M. 2012, *ApJ*, 749, 96
- Hubble, E. P. 1926, *ApJ*, 64, 321
- Ichikawa, S.-I. 1989, *AJ*, 97, 1600
- Jang, I. S., & Lee, M. G. 2014, *ApJ*, 795, L6
- Jedrzejewski, R. I. 1987, *MNRAS*, 226, 747
- Kaufmann, T., Wheeler, C., & Bullock, J. S. 2007, *MNRAS*, 382, 1187
- Kazantzidis, S., Lokas, E. L., Callegari, S., Mayer, L., & Moustakas, L. A. 2011, *ApJ*, 726, 98
- Klypin, A., Karachentsev, I., Makarov, D., & Nasonova, O. 2015, *MNRAS*, 454, 1798
- Klypin, A., Kravtsov, A. V., Valenzuela, O., & Prada, F. 1999, *ApJ*, 522, 82
- Koposov, S., Belokurov, V., Evans, N. W., et al. 2008, *ApJ*, 686, 279
- Lambas, D. G., Maddox, S. J., & Loveday, J. 1992, *MNRAS*, 258, 404
- Lisker, T., Grebel, E. K., Binggeli, B., & Glatt, K. 2007, *ApJ*, 660, 1186
- Lisker, T., Weinmann, S. M., Janz, J., & Meyer, H. T. 2013, *MNRAS*, 432, 1162
- Liu, C., Peng, E. W., Côté, P., et al. 2015, *ApJ*, 812, 34
- Lokas, E. L., Majewski, S. R., Kazantzidis, S., et al. 2012, *ApJ*, 751, 61
- Martin, N. F., de Jong, J. T. A., & Rix, H.-W. 2008, *ApJ*, 684, 1075
- Mastropietro, C., Moore, B., Mayer, L., et al. 2005, *MNRAS*, 364, 607
- Mateo, M. L. 1998, *ARA&A*, 36, 435
- Mayer, L., Governato, F., Colpi, M., et al. 2001, *ApJ*, 547, L123
- Mayer, L., Kazantzidis, S., Mastropietro, C., & Wadsley, J. 2007, *Nature*, 445, 738
- McConnachie, A. W. 2012, *AJ*, 144, 4
- Mei, S., Blakeslee, J. P., Côté, P., et al. 2007, *ApJ*, 655, 144
- Méndez-Abreu, J., Sánchez-Janssen, R., & Aguerri, J. A. L. 2010, *ApJ*, 711, L61
- Méndez-Abreu, J., Sánchez-Janssen, R., Aguerri, J. A. L., Corsini, E. M., & Zarattini, S. 2012, *ApJ*, 761, L6
- Michard, R., & Andreon, S. 2008, *A&A*, 490, 923
- Mieske, S., Hilker, M., Jordán, A., et al. 2008, *A&A*, 487, 921
- Mistani, P. A., Sales, L. V., Pillepich, A., et al. 2016, *MNRAS*, 455, 2323
- Moore, B., Katz, N., Lake, G., Dressler, A., & Oemler, A. 1996, *Nature*, 379, 613
- Moore, B., Lake, G., & Katz, N. 1998, *ApJ*, 495, 139
- Moore, B., Lake, G., Quinn, T., & Stadel, J. 1999, *MNRAS*, 304, 465
- Muñoz, R. P., Puzia, T. H., Lançon, A., et al. 2014, *ApJS*, 210, 4
- Norris, M. A., Kannappan, S. J., Forbes, D. A., et al. 2014, *MNRAS*, 443, 1151
- Okamoto, T., Gao, L., & Theuns, T. 2008, *MNRAS*, 390, 920
- Padilla, N. D., & Strauss, M. A. 2008, *MNRAS*, 388, 1321
- Paudel, S., Duc, P.-A., Côté, P., et al. 2013, *ApJ*, 767, 133
- Peñarrubia, J., Navarro, J. F., & McConnachie, A. W. 2008, *ApJ*, 673, 226
- Peñarrubia, J., Pontzen, A., Walker, M. G., & Koposov, S. E. 2012, *ApJ*, 759, L42
- Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H.-W. 2002, *AJ*, 124, 266
- Penny, S. J., Forbes, D. A., Strader, J., et al. 2014, *MNRAS*, 439, 3808
- Pontzen, A., & Governato, F. 2012, *MNRAS*, 421, 3464
- Raichoor, A., Mei, S., Erben, T., et al. 2014, *ApJ*, 797, 102
- Richardson, J. C., Irwin, M. J., McConnachie, A. W., et al. 2011, *ApJ*, 732, 76
- Rines, K., & Geller, M. J. 2008, *AJ*, 135, 1837
- Roychowdhury, S., Chengalur, J. N., Karachentsev, I. D., & Kaisina, E. I. 2013, *MNRAS*, 436, L104
- Ryden, B. S. 2004, *ApJ*, 601, 214
- . 2006, *ApJ*, 641, 773
- Ryden, B. S., & Terndrup, D. M. 1994, *ApJ*, 425, 43
- Ryś, A., van de Ven, G., & Falcón-Barroso, J. 2014, *MNRAS*, 439, 284
- Saglia, R. P., Bender, R., & Dressler, A. 1993, *A&A*, 279, 75
- Salomon, J.-B., Ibata, R. A., Martin, N. F., & Famaey, B. 2015, *MNRAS*, 450, 1409
- Sánchez-Janssen, R., Aguerri, J. A. L., & Muñoz-Tuñón, C. 2008, *ApJ*, 679, L77
- Sánchez-Janssen, R., Méndez-Abreu, J., & Aguerri, J. A. L. 2010, *MNRAS*, 406, L65
- Sánchez-Janssen, R., Amorín, R., García-Vargas, M., et al. 2013, *A&A*, 554, A20
- Sandage, A., Freeman, K. C., & Stokes, N. R. 1970, *ApJ*, 160, 831
- Secker, J. 1995, *PASP*, 107, 496
- Sérsic, J. L. 1968, *Atlas de galaxias australes (Cordoba, Argentina: Observatorio Astronomico, 1968)*
- Seth, A. C., van den Bosch, R., Mieske, S., et al. 2014, *Nature*, 513, 398
- Smith, R., Davies, J. I., & Nelson, A. H. 2010, *MNRAS*, 405, 1723
- Smith, R., Duc, P. A., Candlish, G. N., et al. 2013a, *MNRAS*, 436, 839
- Smith, R., Sánchez-Janssen, R., Fellhauer, M., et al. 2013b, *MNRAS*, 429, 1066
- Smith, R., Sánchez-Janssen, R., Beasley, M. A., et al. 2015, *MNRAS*, 454, 2502
- Stark, A. A. 1977, *ApJ*, 213, 368
- Starkenburgh, T. K., & Helmi, A. 2015, *A&A*, 575, A59
- Statler, T. S. 1994, *ApJ*, 425, 500
- Sung, E., Han, C., Ryden, B. S., et al. 1998, *ApJ*, 505, 199
- Teyssier, R., Pontzen, A., Dubois, Y., & Read, J. I. 2013, *MNRAS*, 429, 3068
- Toloba, E., Guhathakurta, P., van de Ven, G., et al. 2014, *ApJ*, 783, 120
- Tolstoy, E., Hill, V., & Tosi, M. 2009, *ARA&A*, 47, 371
- Turner, M. L., Côté, P., Ferrarese, L., et al. 2012, *ApJS*, 203, 5
- van den Bergh, S. 2000, *The Galaxies of the Local Group (Cambridge)*

- van den Bosch, R. C. E., & van de Ven, G. 2009, MNRAS, 398, 1117
- van der Wel, A., Rix, H., Holden, B. P., Bell, E. F., & Robaina, A. R. 2009, ApJ, 706, L120
- Vincent, R. A., & Ryden, B. S. 2005, ApJ, 623, 137
- Walker, M. G., Mateo, M., Olszewski, E. W., et al. 2009, ApJ, 704, 1274
- Walker, M. G., & Peñarrubia, J. 2011, ApJ, 742, 20
- Weijmans, A.-M., de Zeeuw, P. T., Emsellem, E., et al. 2014, MNRAS, 444, 3340
- Weisz, D. R., Dolphin, A. E., Skillman, E. D., et al. 2014, ApJ, 789, 147
- Weisz, D. R., Dalcanton, J. J., Williams, B. F., et al. 2011, ApJ, 739, 5
- Wheeler, C., Pace, A. B., Bullock, J. S., et al. 2015, ArXiv e-prints, arXiv:1511.01095
- White, S. D. M., & Rees, M. J. 1978, MNRAS, 183, 341
- Woo, J., Courteau, S., & Dekel, A. 2008, MNRAS, 390, 1453
- Zhang, H.-X., Peng, E. W., Côté, P., et al. 2015, ApJ, 802, 30
- Zolotov, A., Brooks, A. M., Willman, B., et al. 2012, ApJ, 761, 71