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Hubble Frontier Fields: the geometry and dynamics of the massive galaxy cluster merger MACSJ0416.1–2403

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ABSTRACT

We use a joint optical/X-ray analysis to constrain the geometry and history of the ongoing merging event in the massive galaxy cluster MACSJ0416.1–2403 ($z = 0.397$). Our investigation of cluster substructure rests primarily on a combined strong- and weak-lensing mass reconstruction based on the deep, high-resolution images obtained for the Hubble Frontier Fields initiative. To reveal the system's dynamics, we complement this lensing analysis with a study of the intracluster gas using shallow *Chandra* data, and a three-dimensional model of the distribution and motions of cluster galaxies derived from over 100 spectroscopic redshifts. The multiscale grid model obtained from our combined lensing analysis extends the high-precision strong-lensing mass reconstruction recently performed to cluster-centric distances of almost 1 Mpc. Our analysis detects the two well-known mass concentrations in the cluster core. A pronounced offset between collisional and collisionless matter is only observed for the SW cluster component, while excellent alignment is found for the NE cluster. Both the lensing analysis and the distribution of cluster light strongly suggest the presence of a third massive structure, almost 2 arcmin SW of the cluster centre. Since no X-ray emission is detected in this region, we conclude that this structure is non-virialized and speculate that it might be part of a large-scale filament almost aligned with our line of sight. Combining all evidence from the distribution of dark and luminous matter, we propose two alternative scenarios for the trajectories of the components of MACSJ0416.1–2403. Upcoming deep X-ray observations that allow the detection of shock fronts, cold cores, and sloshing gas (all key diagnostics for studies of cluster collisions) will allow us to test, and distinguish between these two scenarios.

Key words: gravitational lensing: strong – gravitational lensing: weak – galaxies: clusters: individual: – dark matter – X-rays: galaxies: clusters.

1 INTRODUCTION

In the course of the past decades, gravitational lensing has become one of the most powerful tools to map the distribution of

dark matter, starting with the confirmation of gravitational lensing as the causal origin of the giant arc in the cluster Abell 370 by Soucail et al. (1988). The bending of light by foreground clusters can be observed in two regimes: the strong-lensing (SL) regime, limited to the densest part of the cluster, i.e. its core, and the weak-lensing (WL) regime, in its outskirts. Gravitational lensing allows

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astronomers not only to directly measure the distribution of the total gravitational mass (dark or luminous), but also to use clusters as ‘cosmic telescopes’ to image very distant galaxies and to constrain the geometry of the Universe (for reviews, see e.g. Massey et al. 2010; Kneib & Natarajan 2011). As the most powerful gravitational telescopes, massive clusters are sought-after observational targets.

Occupying the nodes of the cosmic web of large-scale filaments and sheets (Bond, Kofman & Pogosyan 1996), massive clusters and, specifically, massive cluster mergers are of particular interest also in the context of structure formation studies. The case of the bullet cluster (1E0657–56, $z = 0.3$; Clowe, De Lucia & King 2004) is exceptional in this context: it shows a merging event of two clusters where the merging direction is perpendicular to the line of sight, maximizing the apparent separation and revealing a clear cone-shaped shock front ahead of the smaller merger component. This rare geometry has allowed studies to separately investigate the distribution and dynamics of the baryonic and dark matter components, e.g. by using a combination of strong- and weak-gravitational lensing (Bradač et al. 2006). This approach is particularly powerful as SL constrains precisely the location and shape of the cluster core, while WL maps the mass distribution on larger scales. A similar analysis was performed also on MACSJ0025.4–1222 ($z = 0.58$; Ebeling et al. 2007; Bradač et al. 2008). Although any given observation captures no more than a snapshot of the complex process of cluster growth, the different dynamical behaviour of collisional (gas) and collisionless matter (dark matter and galaxies) often observed in merging systems has been used to great effect to constrain the three-dimensional trajectories of the merger components (e.g. Clowe et al. 2004; Ma, Ebeling & Barrett 2009; Merten et al. 2011; Jauzac et al. 2012; Hsu, Ebeling & Richard 2013). In this paper, we investigate a massive cluster that is both a spectacularly efficient gravitational lens and an active merger.

MACSJ0416.1–2403 ($z = 0.397$, hereafter MACSJ0416) was discovered by the Massive Cluster Survey (MACS; Ebeling, Edge & Henry 2001) and classified as an actively merging system by Mann & Ebeling (2012) based on its X-ray/optical morphology. Because of its large Einstein radius, as revealed in *Hubble Space Telescope* (*HST*) observations obtained for programme GO-11103 (PI: Ebeling), MACSJ0416 was selected as one of five ‘high-magnification’ clusters in the Cluster Lensing And Supernova survey with Hubble (CLASH; Postman et al. 2012). The system’s highly elongated mass distribution, typical of merging systems, allowed numerous strongly lensed galaxies to be discovered (Zitrin et al. 2013; Richard et al. 2014) in these imaging data. More recently, MACSJ0416 was chosen as one of six targets for the Hubble Frontier Fields (HFF) initiative. Launched by the Space Telescope Science Institute (STScI) in 2013, this observing programme aims to harness the gravitational magnification of massive cluster lenses to study the distant Universe to unprecedented depth. The HFF programme allocates 140 *HST* orbits to imaging observations of each cluster, split between three filters on the Advanced Camera for Survey (ACS), and four on the Wide Field Camera 3 (WFC3), to reach an unprecedented depth for cluster studies of $m_{\text{AB}} \sim 29$ in all seven passbands. The HFF observations of MACSJ0416 with ACS, performed in early 2014, allowed us to identify 51 new multiply-imaged galaxies, bringing the total number of lensed images to a record of 194 (Jauzac et al. 2014). The resulting SL mass model confirmed the bimodal mass distribution of MACSJ0416, and constrains the mass within the core region to a precision of better than 1 per cent.

In this paper, we extend the analysis of Jauzac et al. (2014) by using both strong- and weak-gravitational-lensing constraints to

measure and map the mass distribution of MACSJ0416 to larger cluster-centric radii. In addition, we use archival *Chandra* X-ray data as well as radial velocities measured for over 100 cluster galaxies to investigate the relative motions of collisional and collisionless matter in projection, as well as along our line of sight. The result is a model of the three-dimensional geometry and merger history of this complex system. Our paper is organized as follows: observations of MACSJ0416 are summarized in Section 2, an overview of our earlier SL analysis is provided in Section 3, the construction of the WL catalogue is described in Section 4, our gravitational-lensing mass-modelling technique is explained in Section 5, results are presented in Section 6, the dynamical analysis of MACSJ0416 is performed in Section 7, and, finally, a summary is provided in Section 8. We adopt the Λ cold dark matter concordance cosmology with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, and a Hubble constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. Magnitudes are quoted in the AB system.

2 OBSERVATIONS

2.1 Pre-HFF *HST* data

MACSJ0416 was first observed with the *HST* using the Wide Field Planetary Camera 2 in 2007 as part of the SNAPshot programme GO-11103 (PI: Ebeling). This observation established MACSJ0416 as a powerful gravitational lens which led to its inclusion in the CLASH programme (PI: Postman; Postman et al. 2012). Hence, MACSJ0416 was observed with *HST* in 2012 for a total of 20 orbits across 16 passbands, from the UV to the near-IR. Table 1 lists details of these ACS and WFC3 observations, which were used for the pre-HFF analysis of MACSJ0416. All mass models based on pre-HFF data (Coe, Bradley & Zitrin 2014; Johnson et al. 2014; Richard et al. 2014) are publicly available.¹

2.2 Hubble Frontier Fields data

MACSJ0416 is the second cluster to be observed in the HFF programme (GO/DD 13496). Observations with ACS were performed from 2014 January 5 to February 9 in three filters (*F435W*, *F606W*, and *F814W*) for a total exposure time of 20, 12, and 48 orbits, respectively. A summary of these observations is provided at the end of Table 1. At the time of this writing, imaging of the cluster with WFC3 for the HFF programme had yet to commence.

We applied basic data-reduction procedures to the HFF/ACS data, using *HSTCAL* and the most recent calibration files. Individual frames were co-added using *ASTRODRIZZLE* after registration to a common ACS reference image using *TWEAKREG*. After an iterative process, we achieve an alignment accuracy of 0.1 pixel. Our final stacked images have a pixel size of 0.03 arcsec.

2.3 *Chandra* X-ray data

MACSJ0416 was observed with the Advanced CCD Imaging Spectrometer (ACIS-I) on board the *Chandra* X-ray Observatory on 2009-06-07 for 16 ks (ObsID 10446, PI: Ebeling), and on 2014-06-09 for 37 ks (ObsID 16237, PI: Jones). We process these archival data using CIAO v4.6 and CALDB v4.5.9, merging them into a single 53 ks observation. After examining the light curve of the accumulated count rate for periods of enhanced particle background, we extract a raw image of the cluster in the 0.7–7 keV band and use the CIAO tool *mkexpmap* to compute an effective exposure map,

¹ <http://archive.stsci.edu/prepds/frontier/lensmodels/>

Table 1. Summary of *HST* observations of MACSJ0416. The HFF observations are highlighted in bold.

RA (J2000)	Dec. (J2000)	Instrument/Filter	Exposure time (in s)	Programme ID
04 16 07.2	−24 03 35.7	WFC2/F814W	1200	11103
04 16 07.2	−24 03 35.7	WFC2/F606W	1200	11103
04 16 08.4	−24 04 20.0	ACS/F435W	2052	12459
04 16 08.4	−24 04 20.0	ACS/F475W	2064	12459
04 16 08.4	−24 04 20.0	ACS/F606W	2018	12459
04 16 08.4	−24 04 20.0	ACS/F625W	2017	12459
04 16 08.4	−24 04 20.0	ACS/F775W	2031	12459
04 16 08.4	−24 04 20.0	ACS/F814W	4037	12459
04 16 08.4	−24 04 20.0	ACS/F850LP	4086	12459
04 16 08.4	−24 04 20.0	WFC3/F105W	2815	12459
04 16 08.4	−24 04 20.0	WFC3/F110W	2515	12459
04 16 08.4	−24 04 20.0	WFC3/F125W	2515	12459
04 16 08.4	−24 04 20.0	WFC3/F140W	2312	12459
04 16 08.4	−24 04 20.0	WFC3/F160W	5029	12459
04 16 08.4	−24 04 20.0	WFC3/F225W	3634	12459
04 16 08.4	−24 04 20.0	WFC3/F275W	3684	12459
04 16 08.4	−24 04 20.0	WFC3/F336W	2360	12459
04 16 08.4	−24 04 20.0	WFC3/F390W	2407	12459
04 16 08.9	−24 04 28.7	ACS/F435W	52 460	HFF-13496
04 16 08.9	−24 04 28.7	ACS/F606W	31 476	HFF-13496
04 16 08.9	−24 04 28.7	ACS/F814W	125 904	HFF-13496

taking vignetting effects into account. The raw image, which preserves the recorded photon statistics, is adaptively smoothed using `asmooth` (Ebeling, White & Rangarajan 2006), requiring 3σ significance of all features with respect to the local background.

2.4 Spectroscopic and photometric redshifts

More than 100 spectroscopic galaxy redshifts are available within the field of MACSJ0416. Spectroscopic redshifts from Ebeling, Ma & Barrett (2014) are complemented by redshifts obtained for VLT programme 186.A–0798 (Balestra et al., in preparation). We also make use of the catalogue of photometric redshifts derived by the CLASH team from *HST* imaging in 16 passbands (second block of entries in Table 1) utilizing the Bayesian Photometric Redshift programme (Benítez et al. 2004; Coe et al. 2006). We use of all of these redshifts to select background galaxies for our WL catalogue, as well as for the identification (and removal) of cluster members (see Section 4 for details).

3 SL ANALYSIS: REVISITING MULTIPLE IMAGES

Since our HFF SL analysis of MACSJ0416 has already been presented in Jauzac et al. (2014, hereafter J14), we here provide only a brief synopsis of the mass model derived and the main results.

Before the HFF observations of MACSJ0416, Zitrin et al. (2013) identified 23 multiple-image systems (corresponding to a total of 70 individual images) in the CLASH data (Postman et al. 2012). The identifications of 10 of these (comprising 36 individual images) were considered less robust. In Richard et al. (2014, hereafter R14), we included only the most robust systems as well as a few candidate systems showing clear counter images at locations predicted by our preliminary SL analysis. Our final list contained 17 multiply-imaged galaxies (47 individual images). Thanks to the unprecedented depth of the HFF data, J14 were able to dramatically improve these numbers, by discovering 51 new multiple-image systems, bringing our final list of identifications to 68 multiply-imaged galaxies, with a

total of 194 individual images. Spectroscopic confirmation however has so far been obtained for only nine systems. The full list of these systems is given in J14.

Using a subset of the 57 most securely identified multiple-image systems, we built a SL parametric mass model using the publicly available `LENSTOOL`² software. The resulting best-fitting model comprises two cluster-scale dark matter haloes and 98 galaxy-scale haloes. The parameters describing this best-fitting mass model are listed in Table 2. We also provide a mass estimate for both components within 20 arcsec (~ 100 kpc).

This model predicts image positions to within an rms error of 0.68 arcsec, an improvement in precision of almost a factor of 2 over pre-HFF models of this cluster. The total mass enclosed in the multiple-image region is $M_{\text{SL}}(R < 320 \text{ kpc}) = (3.26 \pm 0.03) \times 10^{14} M_{\odot}$. This measurement offers a three-fold improvement in precision and drives the statistical mass uncertainty below 1 percent for the first time in any cluster. Finally, the statistical uncertainty in the median magnification has been lowered to 4 percent. The resulting high-precision magnification map of this powerful cluster lens immediately improves the constraints on the luminosity function of high-redshift galaxies lensed by this system.

For more details on our HFF SL analysis of MACSJ0416, we refer the reader to J14, where the methodology and mass measurements are described in details.

4 WL CONSTRAINTS

In this section, we summarize our analysis methodology and then discuss in particular the enhancements to our technique wrought by significant improvements in the data quality brought about by the HFF observations. A more detailed description of the method used to generate the WL background-galaxy catalogue is presented in (Jauzac et al. 2012, hereafter J12).

² <http://projects.lam.fr/projects/lenstool/wiki>

Table 2. Best-fitting pseudo-isothermal elliptical mass distribution parameters for the two large-scale dark matter haloes. Coordinates are quoted in arcseconds with respect to $\alpha = 64.038\ 1013$ deg, $\delta = -24.067\ 4860$ deg (yellow cross in Fig. 6). Errors correspond to the 1σ confidence level. Parameters in brackets are not optimized. The reference magnitude for scaling relations is $m_{F814W} = 19.8$. Masses are quoted within an aperture of 20 arcsec (~ 100 kpc).

Component	C1	C2	L^* elliptical galaxy
Δ RA	$-4.5^{+0.7}_{-0.6}$	$24.5^{+0.5}_{-0.4}$	–
Δ Dec.	$1.5^{+0.5}_{-0.6}$	$-44.5^{+0.6}_{-0.8}$	–
e	0.7 ± 0.02	0.7 ± 0.02	–
θ	$58.0^{+0.7}_{-1.2}$	37.4 ± 0.4	–
r_{core} (kpc)	$77.8^{+4.1}_{-4.6}$	103.3 ± 4.7	[0.15]
r_{cut} (kpc)	[1000]	[1000]	$29.5^{+7.4}_{-4.3}$
σ (km s $^{-1}$)	779^{+22}_{-20}	955^{+17}_{-22}	147.9 ± 6.2
M ($10^{13} M_{\odot}$)	6.02 ± 0.09	6.12 ± 0.09	–

4.1 The ACS source catalogue

Our WL analysis is based on shape measurements in the ACS/ $F814W$ band. Following a method developed for the analysis of data obtained for the COSMOS survey, and described in Leauthaud et al. (2007, hereafter L07), the SExtractor photometry package (Bertin & Arnouts 1996) is used to detect sources with the ‘Hot–Cold’ method (Rix et al. 2004; J07). It consists of running SExtractor twice: first with a configuration optimized to detect only the brightest objects (the *cold* step), and then a second time with a configuration optimized to detect the faint objects that contain most of the lensing signal (the *hot* step). The resulting catalogue is then cleaned by removing spurious or duplicate detections using a semi-automatic algorithm that defines polygonal masks around stars or saturated pixels. Galaxies are distinguished from stars by examining the distribution of objects in the magnitude (MAG_AUTO) versus peak surface brightness (MU_MAX) plane (see L07 and J12 for more details). Finally, the drizzling process introduces pattern-dependent correlations between neighbouring pixels, which artificially reduces the noise level of co-added drizzled images. Care must be taken to correct for this effect. Because we have used the same Drizzle pixelfrac and convolution kernel parameters as L07, we apply the same remedy as L07 by simply scaling up the noise level in each pixel by the same constant $F_A \approx 0.316$, defined by Casertano et al. (2000). The resulting catalogue comprises 4296 sources identified as galaxies and 1171 sources identified as point sources (stars) within a magnitude limit of $m_{F814W} = 29.5$. Fig. 1 shows the magnitude distribution of the detected galaxies and stars.

Since only galaxies behind the cluster are gravitationally lensed, the presence of cluster members and foreground galaxies dilutes the observed shear and reduces the significance of all quantities derived from it. Identifying and eliminating as many of the contaminating unlensed galaxies is thus crucial. As a first step, we identify cluster galaxies with the help of the catalogue of photometric redshifts compiled by the CLASH collaboration, and the spectroscopic redshifts mentioned in Section 2.4. All galaxies with photometric redshift $0.35 < z_{\text{phot}} < 0.44$ are considered to be cluster galaxies. The spectroscopic cluster membership criterion is defined by

$$z_{\text{cluster}} - dz < z < z_{\text{cluster}} + dz,$$

where z is the spectroscopic redshift of the considered galaxy, $z_{\text{cluster}} = 0.3979$ is the systemic redshift of the cluster, and

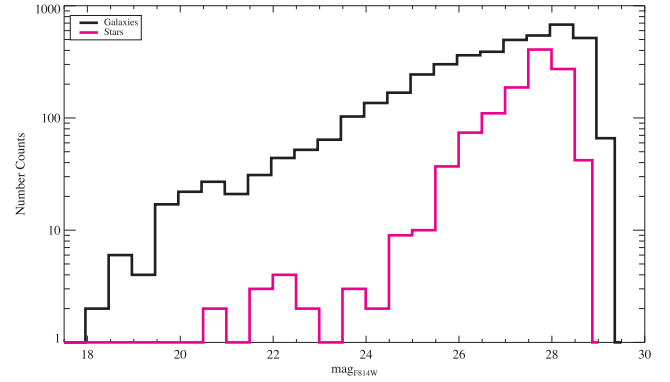


Figure 1. Magnitude distribution of the sources identified as galaxies (solid black) and stars (magenta) resulting from our SExtractor detection.

$dz = 0.0104$ is the 3σ cut defined by the colour–magnitude selections presented in Section 5.2. Only 30 per cent of the sources in our ACS object catalogue have a photometric redshift. Of these 30 per cent, 17 per cent are identified as cluster members or foreground sources following the aforementioned selection criteria. Due to the large difference in depth between the CLASH and HFF observations, the photometric redshift catalogue is not sufficient to identify all the unlensed contaminants in our catalogue. Therefore, taking advantage of the three HFF ACS-band imaging, we use a colour–colour diagram to identify foreground and cluster members (Fig. 2). Using galaxies with photometric or spectroscopic redshifts, we identify the region in colour–colour space that is dominated by unlensed galaxies (foreground galaxies and cluster members) and define its boundaries by $m_{F435W} - m_{F814W} < 0.677\ 76 (m_{F435W} - m_{F606W}) + 0.3$; $m_{F435W} - m_{F814W} > 0.877\ 76 (m_{F435W} - m_{F606W}) - 0.76$; $m_{F435W} - m_{F814W} > 0.3$. We consequently remove all objects within this region from our analysis. Fig. 3 shows the galaxy redshift distribution before and after this $F435W$ – $F606W$ – $F814W$ colour–colour selection. This selection is very efficient at removing cluster members and foreground galaxies at $z \leq 0.44$ – for the subset of our galaxies that have redshifts, 88 per cent of the unlensed population are eliminated.

The final validation of our colour–colour selection is done by predicting the colours expected from spectral templates at the redshift of the cluster or in the foreground. We use empirical templates from Coleman, Wu & Weedman (1980) and Kinney et al. (1996) as well as theoretical templates from Bruzual & Charlot (2003) for various galaxy types in the Hubble sequences (ranging from elliptical to SB) and starburst galaxies. The location of the colour–colour tracks at $z < 0.44$ agree well with our selection region as shown in Fig. 4 for the Bruzual and Charlot model.

4.2 Shape measurements and lensing cuts

Measurements of galaxy shapes are performed using the RRG method (Rhodes, Refregier & Groth 2000). This method has been developed for the analysis of data obtained from space, featuring a small, diffraction-limited point spread function (PSF). It decreases the noise in the shear estimators by correcting each moment of the PSF linearly and only dividing them at the very end to compute an ellipticity. Rhodes et al. (2007) demonstrated that the ACS PSF is not as stable as one might expect from a space-based camera. Its size and ellipticity vary considerably on time-scales of weeks due to telescope ‘breathing’, which induces a deviation from the nominal focus and thus from the nominal PSF which becomes larger

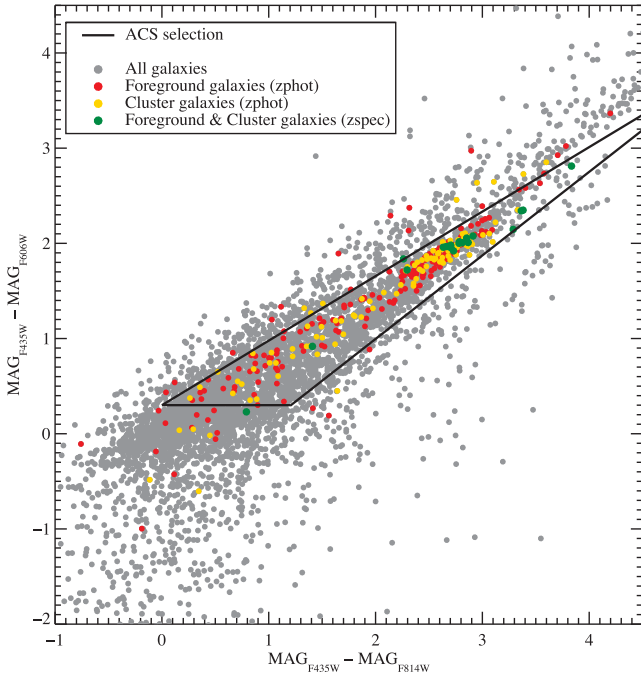


Figure 2. Colour–colour diagram ($m_{F435W} - m_{F814W}$) versus ($m_{F435W} - m_{F606W}$) for objects within the HFF/ACS image of MACSJ0416. Grey dots represent all galaxies in the study area. Unlensed galaxies diluting the shear signal are marked by different colours: galaxies spectroscopically confirmed as cluster members or foreground galaxies (green); galaxies classified as foreground objects because of their photometric redshifts (red); and galaxies classified as cluster members via photometric redshifts (yellow). The solid black lines delineate the colour-cut defined for this work to mitigate shear dilution by unlensed galaxies.

and more elliptical. To overcome this problem, Rhodes et al. (2007) created a grid of simulated PSF images. From a comparison of the ellipticity of ~ 20 stars in each image to the ellipticities of these model images, the effective focus of the observation can be determined. PSF parameters are then interpolated using the method presented in Massey et al. (2002). This technique was used in L07 and J12. However, it has since been shown that PSF variations occur even between subsequent exposures, and thus a modelling of the PSF for each epoch results in a more accurate estimation of the correction to apply to shear estimations (Harvey et al., in preparation).

In order to handle the multi-epoch images of MACSJ0416, we adapted the RRG pipeline (L07) to model the average PSF at the position of each galaxy in the stacked image (see Harvey et al., in preparation for more details). To this end, we first locate the positions of the stars in the reference frame of the drizzled image using SExtractor and both the magnitude–size and magnitude–MU_MAX diagrams. We measure the second- and fourth-order moments of these stars from each exposure and compare them to the TinyTim model for the F814W band. Using the best-fitting TinyTim PSF model, we then interpolate the PSF to the galaxy positions, rotate the moments such that they are in the reference frame of the stacked image, and then take an average over the stack. (Note that our PSF model thus depends on the number of exposures covering a given area and is not necessarily a continuous function across the field.) Since we know the number of exposures that contribute to the image of each galaxy, we can discard shear estimates of galaxies that have fewer than three exposures. Doing so removes all galaxies near the edge of the field and along chip boundaries.

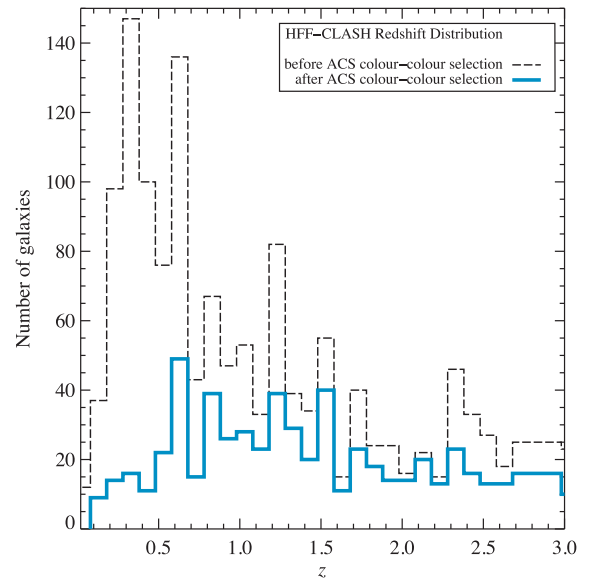


Figure 3. Redshift distribution of all galaxies with m_{F435W} , m_{F606W} , and m_{F814W} photometry from HFF observations, that have photometric or spectroscopic redshifts (dashed black histogram). The cyan histogram shows the redshift distribution of galaxies classified as background objects using the colour–colour criterion illustrated in Fig. 2.

The RRG method returns three parameters: d , a measure of the galaxy size, as well as e_1 and e_2 , the two components of the ellipticity vector $e = (e_1, e_2)$, defined as

$$d = \sqrt{\frac{1}{2}(a^2 + b^2)}$$

$$e = \frac{a^2 - b^2}{a^2 + b^2}$$

$$e_1 = e \cos(2\phi)$$

$$e_2 = e \sin(2\phi).$$

Here, a and b are the major and minor axes of the background galaxy, and ϕ is the orientation angle of the major axis. The ellipticity e is then calibrated by a factor called shear polarisability, G , to obtain the shear estimator $\tilde{\gamma}$:

$$\tilde{\gamma} = C \frac{e}{G}. \quad (1)$$

We use the same global measurement of the shear susceptibility G as in L07:

$$G = 1.125 + 0.04 \arctan \frac{S/N - 17}{4}.$$

Finally, C in equation (1) is the calibration factor, determined using a set of simulated images similar to those used by STEP (Heymans et al. 2006; Massey et al. 2007) for COSMOS images, and is given by $C = (0.86_{-0.05}^{+0.07})^{-1}$ (see L07 for more details).

The last step in constructing the WL catalogue is to exclude galaxies whose shape parameters are so ill-determined that including them would increase the noise in the shear measurements more than the shear signal. These cuts are the same as the ones used in L07 and J12, and are quoted here for clarity.

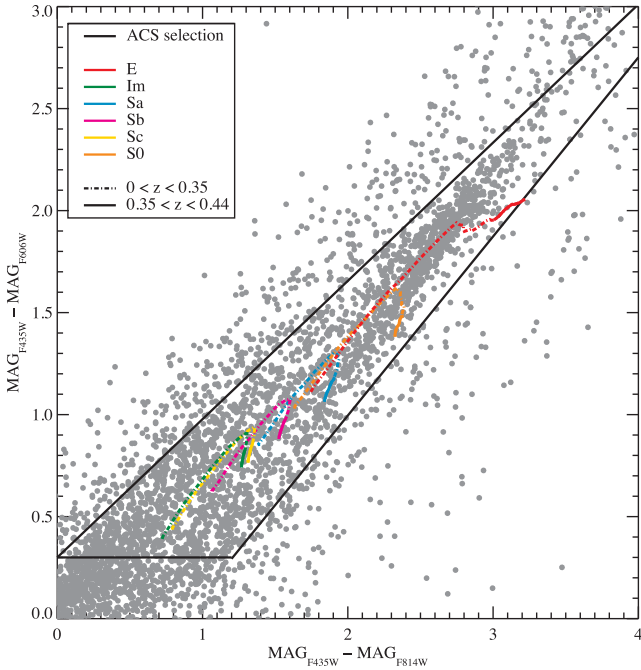


Figure 4. Colour–colour diagram ($\text{mag}_{F435W} - \text{mag}_{F814W}$) versus ($\text{mag}_{F435W} - \text{mag}_{F606W}$) as in Fig. 2. The solid black lines delineate the colour-cut defined for this work. The different spectral templates predicted from Coleman et al. (1980) and Kinney et al. (1996) and the theoretical ones from Bruzual & Charlot (2003) are marked by different colours: elliptical galaxies (red); magellanic irregulars (green); spiral Sa galaxies (cyan); spiral Sb (magenta); spiral Sc (yellow); and S0 galaxies (orange). The dash–dotted curves correspond to a redshift range $0 < z < 0.35$ (foreground galaxies), and the solid ones to $0.35 < z < 0.44$ (cluster members).

(i) Threshold in the estimated detection significance:

$$\frac{S}{N} = \frac{\text{FLUX_AUTO}}{\text{FLUXERR_AUTO}} > 4.5;$$

(ii) threshold in the total ellipticity:

$$e = \sqrt{e_1^2 + e_2^2} < 1;$$

(iii) threshold in the size, as defined by the RRG d parameter:

$$3.6 < d < 30 \text{ pixels.}$$

As explained in J12, the requirement that the galaxy ellipticity be less than unity may appear trivial and superfluous. In practice it is meaningful though since the RRG method allows measured ellipticity values to be greater than 1 due to noise, although ellipticity is by definition restricted to $e \leq 1$. The lower limit in the RRG size parameter d aims to eliminate sources with uncertain shapes, since PSF corrections and thus credible shape measurements become increasingly difficult as the size of a galaxy approaches that of the PSF. The upper limit in d aims to eliminate sources with a size similar to large elliptical cluster members. In addition to applying the aforementioned cuts, and in order to ensure an unbiased mass reconstruction while combining SL and WL, we also remove all background galaxies located in the multiple-image (SL) region, which can be approximated by an ellipse aligned with the cluster elongation as predicted by the SL model ($a = 75$ arcsec, $b = 36$ arcsec, $\theta = 135$ deg, $\alpha = 64.0351$ deg, $\delta = -24.0745$ deg).

Fig. 5 compares the magnitude distribution of selected background galaxies for the pre-HFF data (R14) and the HFF data.

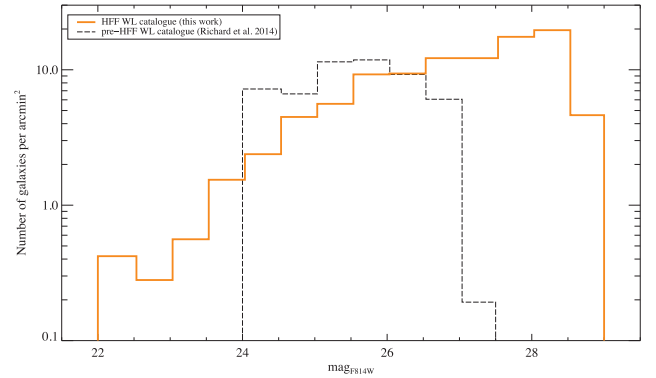


Figure 5. Magnitude distribution of background galaxies selected for our pre-HFF analysis per arcmin² (dashed black line; R14), and for this work (orange).

The HFF-based catalogue extends to ACS- $F814W$ magnitudes of 29, two magnitudes fainter than the pre-HFF data set. Note that the shown distributions differ also at lower magnitudes. Owing to the greatly increased depth of the HFF data compared to those obtained for the CLASH programme, the contamination by faint foreground and cluster galaxies is much increased too. As a consequence, our colour–colour selection is more drastic and removes more objects in the magnitude range $24 < m_{F814W} < 26.13$ galaxies with $m_{F814W} < 24$ are still included in our HFF catalogue that were removed from the pre-HFF catalogue of R14. The reason is the different colour–colour selection employed by R14, which was less efficient and required a magnitude cut at $m_{F814W} = 24$ to remove bright objects. The depth of the HFF images also causes more stars to be saturated, requiring the size of the corresponding masks to be increased; in total ~ 40 per cent of the ACS surface is masked out as a result. Our final WL catalogue is composed of 714 background galaxies, corresponding to a density of ~ 100 galaxies arcmin⁻². Compared to the catalogue generated by our pre-HFF analysis (R14), the density of weakly lensed galaxies has almost doubled.

5 MASS MODELLING: COMBINING SL AND WL CONSTRAINTS

As, in this study, we aim to detect dark-matter-dominated structures outside the SL region, we add a grid-based model to the parametric lens model described in Section 3.

5.1 Grid method: combining SL and WL constraints

In previous work, we modelled the matter distribution with a set of radial basis functions (RBFs) located at the nodes of a multiscale grid, which covered an area slightly larger than observed (Jullo & Kneib 2009; J12; Jullo et al. 2014). We complemented this grid-based model with dPIE potentials (Elíasdóttir et al. 2007) to account for the lensing contribution of cluster members.

In this work, we adopt a slightly different approach. We keep the parametric model described in Section 3 fixed at the best-fitting values and then estimate the RBF amplitudes from the WL constraints. By doing so, we appropriately weigh the SL constraints and do not account for them twice. Indeed, an attempt at optimizing the RBF amplitudes with both SL and WL constraints failed to produce

physically meaningful results as the SL signal completely dominated the optimization process and essentially overwhelmed the WL data. The parametric model contains two cluster-scale haloes and 146 galaxy-scale haloes, as we extend our analysis to the full ACS field of view (see next subsection). We add a uniform grid of RBFs to these main mass components. Each RBF is axisymmetric, fixed in position and size, and only its amplitude varies, as if they were pixels in an image. The radial profile of each RBF is modelled with a dPIE potential (Elíasdóttir et al. 2009). The core radius s is set to the distance between an RBF and its closest neighbour, and the cut radius t is assumed to be three times the core radius (Jullo & Kneib 2009).

We tried different prescriptions for the grid resolution. The optimum solution was achieved using a uniform grid with 2741 RBFs separated from each other by $s = 5.5$ arcsec (see Section 6.1.1 for more details). In addition, we found it necessary to remove the RBFs in the central SL area since, due to a lack of WL constraints, the reconstruction in this region was very noisy. The reconstruction is smooth, thanks to overlapping RBFs.

We sum the contribution of the two components of our model to the observed ellipticity as follows

$$\mathbf{e}_m = M_{\gamma v} \mathbf{v} + \mathbf{e}_{\text{param}} + \mathbf{n}. \quad (2)$$

Here the vector \mathbf{v} contains the amplitudes of the 2741 RBFs, vector $\mathbf{e}_m = [\mathbf{e}_1, \mathbf{e}_2]$ contains the individual shape measurements of the WL sources, $\mathbf{e}_{\text{param}}$ is the fixed ellipticity contribution from the parametric model. The vector \mathbf{n} represents the Gaussian noise in the shape measurements, i.e. the intrinsic ellipticity of galaxies and the noise in our measurements of their shapes. The transformation matrix $M_{\gamma v}$ contains the cross-contribution of each individual RBF to each individual WL source. Each shear component is scaled by the ratio of the distances between each individual source S , the cluster L , and the observer O . The elements of the $M_{\gamma v}$ matrix for the two shear components are

$$\Delta_1^{(j,i)} = \frac{D_{LSi}}{D_{OSi}} \Gamma_1^i(|\theta_i - \theta_j|, s_i, t_i), \quad (3)$$

$$\Delta_2^{(j,i)} = \frac{D_{LSi}}{D_{OSi}} \Gamma_2^i(|\theta_i - \theta_j|, s_i, t_i), \quad (4)$$

where analytical expressions for Γ_1 and Γ_2 are given in Elíasdóttir et al. (2009, equation A8). Note that the shear in the cluster can be large, and the assumption shown in equation (3) may not be strictly valid. However, since the parametric model accounts for most of the lensing effect, the contribution to the grid-based model originates primarily in the WL regime.

5.2 Modelling of cluster members

Complementing our grid of RBFs, we add the contributions from 146 cluster member galaxies (presented in Section 5.1), modelled again as dPIE potentials and selected following the method presented in R14. We define cluster members to be those galaxies that fall within 3σ of a linear model of the cluster red sequence in both the $(m_{F606W} - m_{F814W})$ versus m_{F814W} and the $(m_{F435W} - m_{F606W})$ versus m_{F814W} colour–magnitude diagrams. The magnitude limit of this sample is $m_{F814W} = 23.4$.

These galaxies are then inserted in the model as small-scale perturbators. Their cut radius and velocity dispersions are fixed, and scaled from their luminosities in *HST*/ACS *F814W* band. We derive L^* in our filter of observation based on the K^* magnitudes obtained by Lin et al. (2006) as a function of cluster redshift. Cut radius and

velocity dispersion are then scaled relative to an $m^* = 19.76$ galaxy with velocity dispersion $\sigma^* = (119 \pm 20)$ km s⁻¹ and cut radius $r_{\text{cut}}^* = (85 \pm 20)$ kpc.

5.3 Priors and MCMC sampling

We sample the huge parameter space with the MassInf algorithm implemented in the Bayesys library (Skilling 1998), itself implemented in LENSTOOL, and described in Jullo et al. (2014). In summary, this algorithm uses the Gibbs sampling approach in which, at each iteration, the most significant RBFs are first identified and then adjusted in amplitude to fit the ellipticity measurements. The number of significant RBFs is a prior of MassInf, although we have shown in Jullo et al. (2014) that it has little impact on the reconstruction. We set the initial number of significant RBFs to 2 per cent, and the algorithm converged to about 4 per cent. In contrast to previous analyses, we here do not assume that the resulting mass distribution has to be positive everywhere. We found that incorporating such a prior introduces a spurious bias favouring positive values of the mass–sheet degeneracy.

The objective function is a standard likelihood function, in which noise is assumed to be Gaussian. The algorithm returns a large number of Monte Carlo Markov Chain (MCMC) samples, from which we can estimate mean values and errors on different quantities (mass density field, amplification field, etc.).

5.4 Redshift estimation for background sources

Of the 714 background galaxies in our catalogue, 236 have a photometric redshift estimated from the CLASH data that allow us to isolate background galaxies. We found the following function to provide a good description of the redshift distribution of these background galaxies:

$$\mathcal{N}(z) \propto e^{-(z/z_0)^\beta} \quad (5)$$

with $\beta = 1.84$ and a median redshift $\langle z \rangle = 1.586 = 0.56 z_0$ (Natarajan & Kneib 1997; Gilmore & Natarajan 2009).

In addition, we split the catalogue into a bright and a faint subsample at the median magnitude $m_{F814W} = 26.4$. Within the uncertainties given by the number statistics, the resulting two histograms have the same slope. Since LENSTOOL allows each source to have its own redshift, we randomly draw (during the initialization phase) redshifts from the fitted redshift distribution for all galaxies without spectroscopic or photometric redshift.

6 RESULTS

We now present our results for the properties of both dark and luminous matter in MACSJ0416, beginning with the distribution of the total gravitating mass as reconstructed by our lensing analysis. All masses for the cluster, as well as density profiles, are measured with respect to the position of the brightest cluster galaxy (BCG), i.e. $\alpha = 04:16:09.144$, $\delta = -24:04:02.95$.

6.1 Distribution of total mass

6.1.1 Grid resolution

To assess whether the grid of RBFs presented in Section 5 is optimally suited to describe MACSJ0416, we compared the Bayesian

evidence resulting from the optimization for grids of different resolution. We remind the reader that the logarithmic Bayesian evidence is given by

$$\log(E) = \int_0^1 \langle \log(L) \rangle^\lambda d\lambda,$$

where the average is computed over a set of 10 MCMC realizations at any given iteration step λ , and the integration is performed over all iterations λ_i from the initial model ($\lambda = 0$) to the best-fitting result ($\lambda = 1$). The increment $d\lambda$ depends on the variance between the 10 likelihoods computed at a given iteration, and on a convergence rate that we set equal to 0.1 (see Bayesys manual for details). The increment gets larger as the algorithm converges towards 1.

To test the impact of higher resolution, we created a uniform grid containing 3883 RBFs separated by $s = 3.3$ arcsec. We found the resulting model added noise where no structures are detected in the optical. The Bayesian evidence obtained for this high-resolution grid is $\log(E) = -267$, compared to $\log(E) = -251$ for the grid described in Section 5. We also explored a low-resolution grid, containing 722 RBFs separated by $s = 11.1$ arcsec. The optimization using this grid resulted in a better Bayesian evidence of $\log(E) = -229$, and a similar χ^2 . However, all structures were poorly resolved. Since, in addition, the signal-to-noise ratios in the substructures were equivalent, we discarded this low-resolution model.

6.1.2 Comparison with previous analyses

Using the grid method described in Section 5, we reconstruct the mass distribution of MACSJ0416 within the HFF ACS field of view. Fig. 6 shows contours (white) of the resulting surface mass density overlaid on the ACS colour image. The mass distribution is highly elongated and features no pronounced substructure at significant distances from the major axis, which is typical of merging clusters. We note as well that the mass distribution is more elliptical than with the SL model. This means that the initial SL model was not only under predicting the mass density, but also the shear in the outskirts. Mirroring the system's morphology in the optical and X-ray regime (magenta and cyan contours in Fig. 6) our lensing reconstruction of the total gravitational mass yields again a strongly bimodal distribution. In order to measure the radial mass profile, we define the global cluster centre at $\alpha = 64.0364$ deg and $\delta = -24.0718$ deg (this is the same centre as used in J14 and marked by a yellow cross in Fig. 6). As a test of the consistency of our mass reconstruction techniques, we can compare the projected mass of $M_{\text{SL}}(R < 320 \text{ kpc}) = (3.26 \pm 0.03) \times 10^{14} h_{70}^{-1} M_{\odot}$ measured by J14 using exclusively SL features with the value of $M_{\text{SL+WL}}(R < 320 \text{ kpc}) = (3.15 \pm 0.13) \times 10^{14} h_{70}^{-1} M_{\odot}$ obtained by our joint SL and WL analysis. The precision attained here is lower since, at this distance from the cluster core, WL constraints already contribute to the total mass. Both methods also yield very similar maps of the mass distribution within the cluster core region.

Fig. 7 compares the surface density profiles obtained by this analysis with the ones reported by J14 based on their SL mass model. Note the very good agreement at the outer edge of the multiple-image region (shaded area). At larger radii ($R > 250 h_{70}^{-1} \text{ kpc}$) the predictive power of the model of J14 weakens, due to the lack of structure massive enough to induce SL. By contrast, the combined SL and WL mass model remains sensitive to the less pronounced density variations at the outskirts of the cluster.

Pre-HFF SL+WL mass models were also published in Gruen et al. (2014), and R14. Gruen et al. (2014) present a WL

analysis of the Wide-Field Imager SZ Cluster of Galaxy sample based on data collected with the Wide-Field Imager (Baade et al. 1999) on the 2.2 m MPG/ESO telescope at La Silla. For MACSJ0416, Gruen et al. (2014) obtain mass values that agree well with ours: $M(R < 200 \text{ kpc}) = 1.40_{-0.23}^{+0.22} \times 10^{14} h_{70}^{-1} M_{\odot}$ (Gruen, private communication) versus our value of $M(R < 200 \text{ kpc}) = (1.66 \pm 0.05) \times 10^{14} h_{70}^{-1} M_{\odot}$. More recently, Umetsu et al. (2014) published results from a WL analysis of the CLASH cluster sample using Subaru data. For MACSJ0416, Umetsu et al. (2014) obtain a mass value of $M(R < 950 \text{ kpc}) = 0.98 \pm 0.14 \times 10^{15} h_{70}^{-1} M_{\odot}$ (Umetsu, private communication) that agrees well with our measurement of $M(R < 950 \text{ kpc}) = 1.15 \pm 0.07 \times 10^{15} h_{70}^{-1} M_{\odot}$. Umetsu et al. (2014) reconstructed projected mass density profiles for their cluster sample from a joint likelihood analysis of Subaru shear and magnification measurements. The errors include the estimated contribution from uncorrelated large-scale structures projected along the line of sight. Their result is shown in Fig. 8 and found to be in good agreement with the density profile obtained by us for MACSJ0416. Finally, the pre-HFF model presented by our team in R14, which combines SL and WL constraints but uses a parametric approach similar to the one presented in J14, yields $M(R < 200 \text{ kpc}) = (1.63 \pm 0.03) \times 10^{14} h_{70}^{-1} M_{\odot}$.

6.1.3 Detection of substructure

As mentioned previously, the SL-only results from J14 are in excellent agreement with those presented here from an analysis that combines SL and WL. The advantages of the latter come to bear particularly in the outskirts of the region probed by the HFF images, where any sufficiently massive substructures will reveal their presence by creating weak gravitational shear in the images of background galaxies.

In the case of MACSJ0416, we detect two new substructures at high significance in our mass map; their centres are labelled S1 and S2 in Fig. 6. The projected masses of S1 and S2, estimated within an aperture of ~ 100 kpc, are (4.22 ± 0.56) and $(1.5 \pm 0.20) \times 10^{13} h_{70}^{-1} M_{\odot}$, respectively, values typical of galaxy groups. A tentative identification of S1 as a group of galaxies is supported by the presence of a coinciding galaxy overdensity, clearly visible also as a peak in the cluster light distribution (Fig. 6). For S2, the association with a galaxy group is much less obvious. Physical characteristics of S1 and S2 are listed in Table 3. The distribution of cluster light as shown in Fig. 6 is obtained by smoothing a map of the flux in the ACS-F814W band from cluster members with a Gaussian kernel ($\sigma = 9$ arcsec).

The imprint of these minor mass concentrations on the overall mass density profile can be seen in Fig. 7 in the form of enhancements at $\sim 450 h_{70}^{-1} \text{ kpc}$ (S2), and (much more clearly) at $\sim 650 h_{70}^{-1} \text{ kpc}$ (S1) from the cluster centre. In order to test whether these features are indeed caused by S1 and S2, we follow the same procedure as used in J12 to confirm the detection of the large-scale filament in the field of MACSJ0717.5+3745 and define three triangular regions, anchored at the global cluster centre and extending towards S1, S2, and (for control purposes) a region towards the NW of the field centre that is void of any mass overdensities. Fig. 9 presents the resulting surface density profiles in these three regions and shows indeed that the excess surface mass density can be attributed to S1 and S2, while no significant variations are observed in the radial surface mass density profile of the control field.

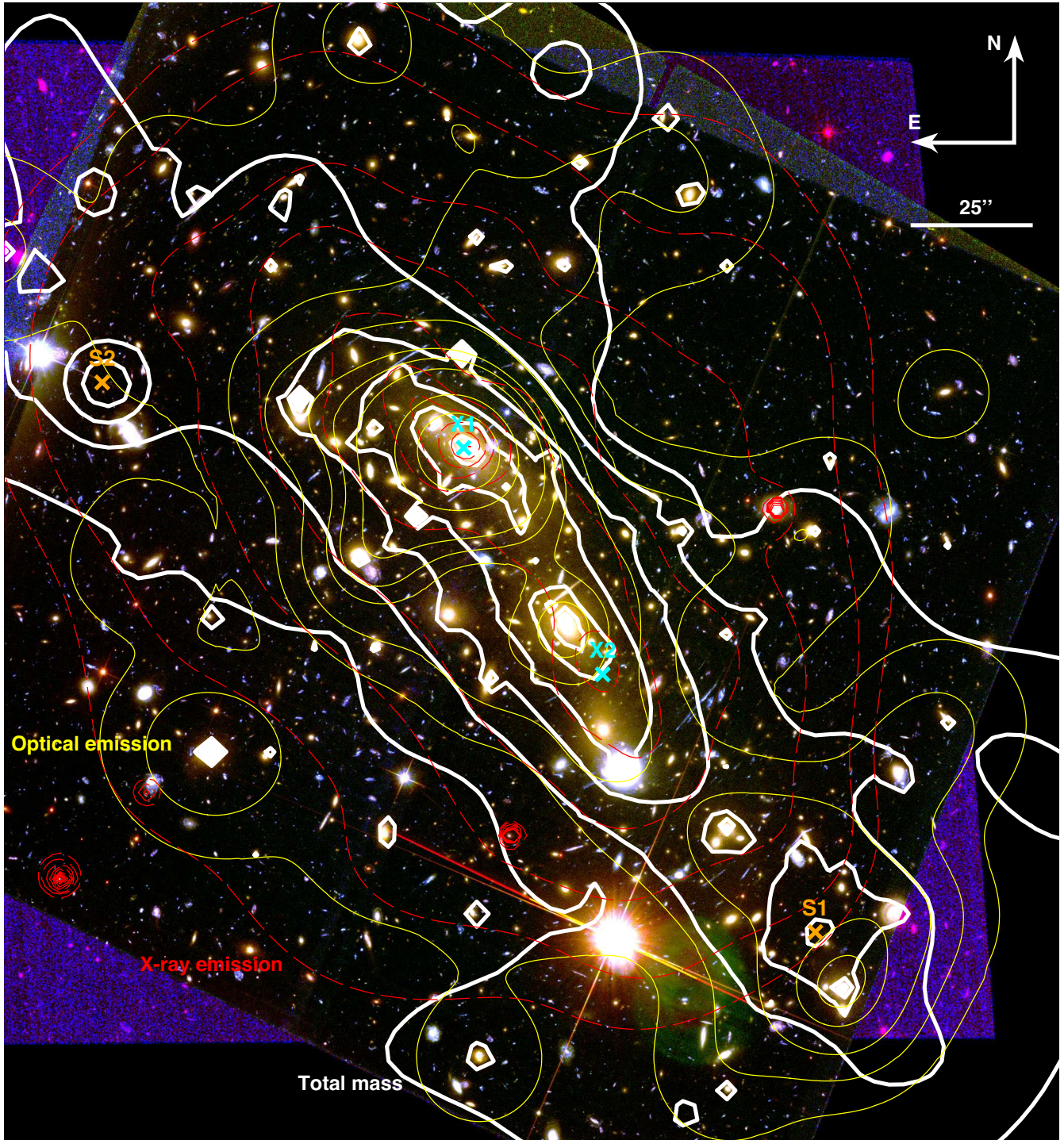


Figure 6. Composite colour image of the galaxy cluster MACSJ0416 created from *HST*/ACS images in the *F814W*, *F606W*, and *F435W* passbands. Mass contours from our gravitational-lensing analysis are shown in bold white, contours of the adaptively smoothed X-ray surface brightness in the 0.5–7 keV band as observed with *Chandra* are shown in dashed red, while the light distribution is delineated by yellow contours. Substructures S1 and S2 are marked with orange crosses while the two X-ray peaks, labelled X1 and X2, are marked by cyan crosses.

6.2 Distribution of stellar mass

To measure the stellar mass distribution, M_* , across the ACS field of view, we use the same method as in J12. We compute the relation $\log(M_*/L_K) = az + b$, established by Arnouts et al. (2007) for quiescent (red) galaxies in the VVDS sample (Le Fèvre et al. 2005), and adopt a Salpeter initial mass function (IMF). Here, L_K

is the galaxy’s luminosity in the K band, z is its redshift, and the parameters a and b are given by

$$a = -0.18 \pm 0.03,$$

$$b = -0.05 \pm 0.03.$$

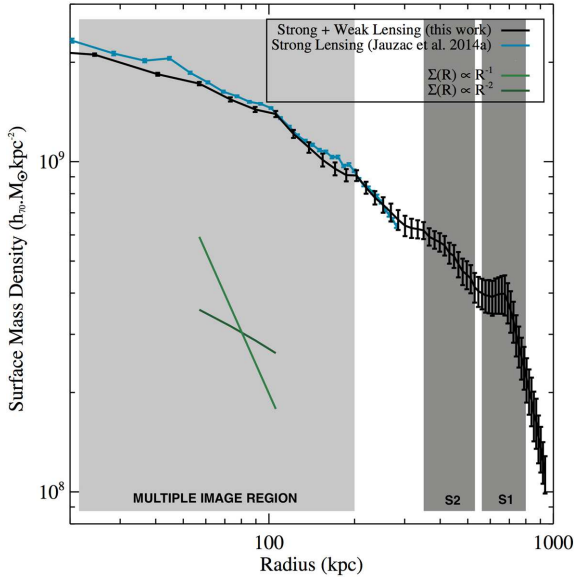


Figure 7. Surface density profile obtained with our comprehensive gravitational-lensing analysis (black). For comparison, we show the surface density profile obtained by J14 based exclusively on SL features (cyan). The shaded light grey area marks the region within which multiple-image systems were found. The two dark grey shaded areas mark the region of substructures S1 and S2. The two green lines represent two different slopes: R^{-1} and R^{-2} .

We apply this relation to our catalogue of 146 cluster members used in our mass model (see Section 5.2). We estimate the K -band luminosity of cluster members observed in the $F814W$ band using theoretical models from Bruzual & Charlot (2003) to predict the typical $(m_{F814W} - m_K)$ colours. We assume a passively evolved galaxy observed at $z = 0.4$, with a range of exponentially decaying star formation histories within the range $\tau = 0.1$ – 2 Gyr. This provides a typical colour $m_{F814W} - m_K = 1.14 \pm 0.04$ (AB system). Using the public data obtained from GSAOI observations in the K_s band (Schirmer et al. 2014), we confirmed this colour for cluster member galaxies located in the central 100×100 arcmin² of the cluster. The resulting projected mass density in stars decreases in proportion to the total projected mass density depicted in Fig. 7. We measure a mass-to-light ratio across the study area of $M_*/L_K = 0.99 \pm 0.03 M_\odot/L_\odot$. To compare our results with those obtained by Leauthaud et al. (2012) for COSMOS data, we need to adjust our measurements to account for the different IMF used by these authors. Applying a shift of 0.25 dex to our masses to convert from a Salpeter IMF to a Chabrier IMF, we find $(M_*/L_K)_{\text{Chabrier}} = 0.78 \pm 0.02 M_\odot/L_\odot$ for quiescent galaxies at $z \sim 0.4$, in good agreement with Leauthaud et al. (2012).

The fraction of the total mass in stars, f_* , i.e. the ratio between the stellar mass and the total mass of the cluster derived from our lensing analysis within the ACS field of view, is $f_* = 3.15 \pm 0.57$ per cent (assuming a Salpeter IMF). The latter value is slightly higher than the one derived by Leauthaud et al. (2012). The difference might be due to different limiting K -band magnitudes or differences in the galaxy environments probed (the COSMOS study was conducted for groups with a halo masses between 10^{11} and $10^{14} h_{70}^{-1} M_\odot$, and extrapolated to haloes of $\sim 10^{15} h_{70}^{-1} M_\odot$). Another cause might be the use of the analytical Arnouts et al. (2007) relation to estimate the stellar mass. As discussed by Ilbert et al. (2010), the M_*/L_K relation used here is only calibrated for massive galaxies, while in practice this ratio varies with galaxy age and colour. Therefore, the Arnouts

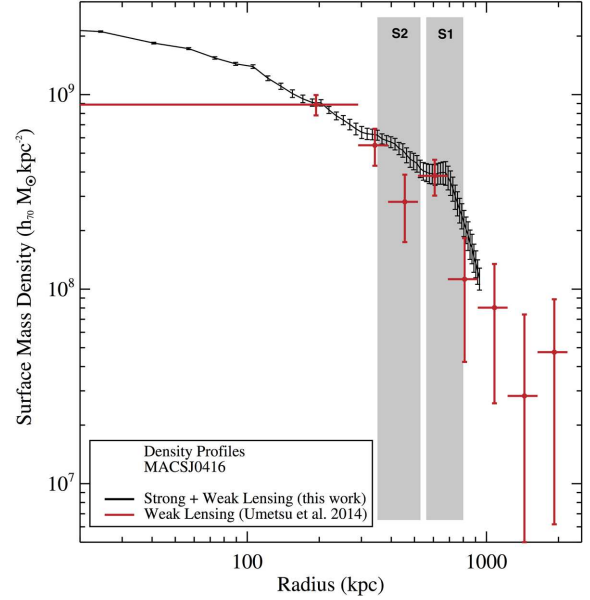


Figure 8. Surface density profile obtained with our comprehensive gravitational-lensing analysis (black). For comparison, we show the surface density profile obtained by Umetsu et al. (2014) based exclusively on WL features (red) derived thanks to Subaru observations (Umetsu, private communication).

et al. (2007) relation overestimates the stellar masses of low-mass galaxies. Although we do not expect our cluster member sample to be dominated by low-mass galaxies, a bias cannot be firmly ruled out. Finally, we compute the total stellar mass within our study area and find $M_* = (3.10 \pm 0.01) \times 10^{13} h_{70}^{-1} M_\odot$. Upcoming HFF data in $F160W$ will provide more direct estimates of the stellar masses.

6.3 Intracluster medium

6.3.1 X-ray morphology

Fig. 6 shows the X-ray contours (in cyan) of the adaptively smoothed X-ray emission as observed with *Chandra*, as described in Section 2.3. The X-ray emission shows two peaks (labelled X1 and X2 in Fig. 6), located at RA = 64.038 458 deg, Dec. = -24.067 361 deg (X1) and RA = 64.029 792 deg, Dec. = -24.080 25 deg (X2), and exhibits a strong elongation in the NE–SW direction. The main peak of the X-ray emission (X1) coincides with the first mass concentration detected in the cluster core by our lensing analysis (C1 in Table 2). However, the second peak (X2) located ~ 45 arcsec (~ 250 kpc) SW of X1, does not coincide with the mass concentration C2 in Table 2 and is offsetted to the SW by ~ 15 arcsec. In the following, we use several models to characterize the gas distribution that gives rise to the X-ray morphology of MACSJ0416.

Acknowledging that the X-ray emission from MACSJ0416 is clearly neither unimodal nor spherically symmetric, we attempt to model the observed X-ray surface brightness distribution as the superposition of two elliptical β -models in *SHERPA*, leaving the position of the centroids free to vary. This model provides a good description of the data and yields best-fitting centroids for the two components of RA = 64.040 604 deg, Dec. = -24.066 54 deg and RA = 64.029 713 deg, Dec. = -24.081 072 deg, respectively. The 1σ uncertainty of these positions is approximately 2 arcsec. The model returns best-fitting values of $r_{c1} = 152 \pm 24$ kpc and $r_{c2} = 68 \pm 17$ kpc for the core radii of C1 and C2, respectively.

Table 3. Coordinates, mass within an ~ 100 kpc aperture, significance of detection, and distance to the cluster centre (D_{C-S}) for the two substructures detected in the outskirts of MACSJ0416.

ID	RA (deg)	Dec. (deg)	$M(10^{13} h_{70}^{-1} M_{\odot})$	σ	$D_{C-S}(\text{kpc})$
S1	64.016 542	-24.094 906	4.22 ± 0.56	7.5	580
S2	64.060 97	-24.063 636	1.46 ± 0.20	7.3	470

The centroid of the main component is thus slightly shifted (by ~ 12 arcsec) to the NE of the X-ray peak, which is not surprising given the irregular morphology of the X-ray emission. The centroid of the second model component however coincides with the X-ray surface-brightness peak X2.

Proceeding to less massive structures identified in our reconstruction of the mass distribution in MACSJ0416 (Fig. 6), we also search for X-ray emission from substructure S1, tentatively identified as a galaxy group in Section 6.1.3. No evidence of X-ray emission from S1 is discernible in Fig. 6. In order to obtain a quantitative assessment of the X-ray luminosity and thus gas mass of S1, we add a third, group-sized component ($r_c = 100$ kpc, $\beta = 0.7$) to our model, tied to the position of S1 as defined in Table 3. We find that the data do not require this third component. The upper limit to the 0.7–7 keV photon flux of $3.2 \times 10^{-6} \text{ cm}^{-2} \text{ s}^{-1}$ within a 20 arcsec radius (90 percent confidence level) corresponds to an upper limit to the X-ray luminosity and gas mass of mass concentration S1 of $L_X < 6.2 \times 10^{42} \text{ erg s}^{-1}$ (unabsorbed, 0.1–2.4 keV) and $M_{\text{gas}}(< 20 \text{ arcsec}) < 2.9 \times 10^{11} h_{70}^{-1} M_{\odot}$, respectively.

6.3.2 Spectral X-ray analysis

In order to further constrain fundamental properties of the intra-cluster medium (ICM) of MACSJ0416, we examined the X-ray spectrum of the diffuse cluster emission within a radius of 2 arcmin from the primary X-ray peak (see Fig. 10). Obvious point sources were excised from the event file prior to the spectral extraction, and a source-free region was defined within the same ACIS-I chip to estimate the local background. We modelled the spectrum with a single-temperature APEC model (Smith et al. 2001) absorbed by the Galactic hydrogen column density, which we fixed at the 21 cm value of $N_{\text{H}} = 3.05 \times 10^{20} \text{ cm}^{-2}$ (Kalberla et al. 2005). Because of the poor photon statistics of the archival ACIS-I observations (~ 4500 source counts in the 0.7–7 keV band), we grouped the spectral channels to obtain a minimum of 20 counts per bin and used the C-statistic (Cash 1979) for the fitting procedure. We obtained an average cluster temperature of $11.0^{+1.4}_{-1.3}$ keV and an Fe abundance of $0.20^{+0.09}_{-0.08} Z_{\odot}$, where the quoted uncertainties represent the 1σ confidence level. According to the M – T relation of Arnaud, Pointecouteau & Pratt (2005), this temperature corresponds to a total mass of $M_{2500} = (4.8^{+0.9}_{-0.7}) \times 10^{14} h_{70}^{-1} M_{\odot}$. With R_{2500} corresponding roughly to 400 kpc, this result is consistent with the value of $M(R < 400 \text{ kpc}) = (4.12 \pm 0.17) \times 10^{14} h_{70}^{-1} M_{\odot}$ measured by our lensing analysis. Most studies (e.g. Nagai, Vikhlinin & Kravtsov 2007; Nelson et al. 2012; Rasia et al. 2012) predict that the lack of thermalization of the gas in violent cluster mergers should lead to an underestimation of the X-ray mass compared to the true mass. However, according to our analysis MACSJ0416 lies on the M – T relation, in spite of ongoing merger activity. While not of great significance in its own right, this result agrees with the low scatter observed around that relation in cluster samples (e.g. Mahdavi et al. 2013). This is important for future X-ray surveys (e.g. *eROSITA*), which will use the X-ray temperature as a proxy for cluster mass.

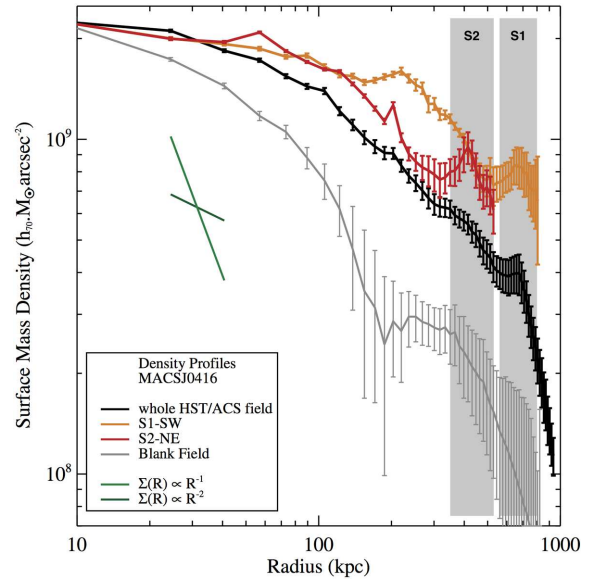


Figure 9. Surface density profiles obtained with our complete gravitational-lensing analysis (black curve). In orange we show the density profile we obtain in a triangular region designed with the cluster centre and including S1. We show the same density profile in red but for S2. Finally, we designed a triangular region into which there is no apparent substructures, and it is shown in grey. The two dark grey shaded areas mark the region of substructures S1 and S2.

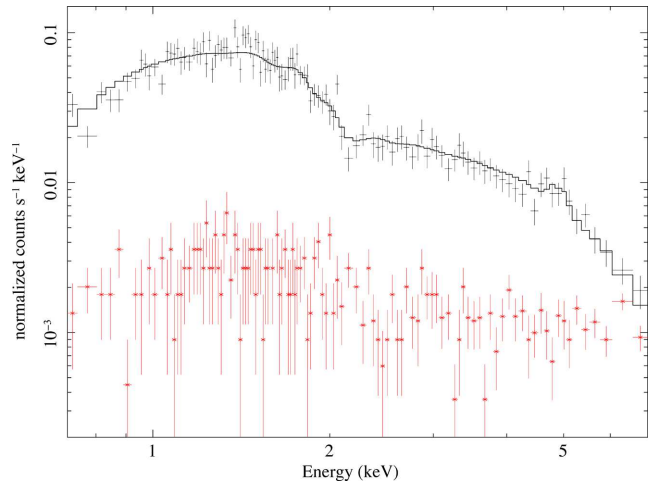


Figure 10. Global *Chandra*/ACIS-I spectrum of the cluster within 2 arcmin radius around the X-ray peak. The solid line shows the best-fitting single-temperature APEC model. The red data points indicate the background level as estimated from a source-free region.

A region of special interest is the core of the NE cluster which appears very compact in Fig. 6 and perfectly aligned with the associated BCG. To test the hypothesis that this cluster component might host a cool core, we extracted the X-ray spectrum of the two components of MACSJ0416 separately. We measure $kT = 10.3^{+1.1}_{-0.8}$ keV and $kT = 13.6^{+2.2}_{-1.9}$ keV for the NE and SW subcluster, respectively. Attempts to directly fit an isothermal plasma model to the current archival data within a circle of 10 arcsec radius of X1 yield unphysical results of either extreme excess absorption of several 10^{21} cm^{-2} (equivalent column density of neutral hydrogen) or temperatures well over 20 keV, i.e. far outside the range that can be constrained with *Chandra*. Given the poor photon statistics (less

than 700 net photons), we do not take these results at face value but rather as indication that an isothermal model is inappropriate. Although we have thus currently no direct spectroscopic evidence of a cool core of the NE cluster component, the data appear to suggest the presence of multiphase gas in this region.

6.3.3 Gas density and gas mass

We follow the procedure described in Eckert et al. (2012) to estimate the three-dimensional gas-density and gas-mass profiles of MACSJ0416. Accounting for vignetting effects, we extracted a surface-brightness profile for a set of concentric annuli of 5 arcsec width centred on the primary X-ray peak, and estimated the local background at radii beyond 4 arcmin. Cluster emission is detected out to ~ 3 arcmin (~ 1 Mpc). The resulting profile was deprojected assuming spherical symmetry using the method of Kriss, Cioffi & Canizares (1983). We converted the deprojected profile into an emission-measure profile assuming a constant temperature of 7.8 keV (see above), and inferred the gas-density profile by assuming constant density within each radial shell. Finally, the gas-mass profile was calculated by integrating the gas-density profile in concentric shells.

We measure gas masses of $M_{\text{gas}}(R < 500 \text{ kpc}) = (3.4 \pm 0.2) \times 10^{13} h_{70}^{-1} M_{\odot}$ and $M_{\text{gas}}(R < 1 \text{ Mpc}) = (8.6 \pm 0.7) \times 10^{13} h_{70}^{-1} M_{\odot}$. In this context, a note is in order regarding systematic effects. Given the irregular morphology of the cluster, the assumption of spherical symmetry might lead to an incorrect gas mass. However, since the morphology of the system beyond the inner regions appears relatively regular, only the gas masses observed in the central regions are significantly affected. Moreover, as stated in Rasia et al. (2011), gas-mass measurements are relatively unaffected by the presence of merging substructures, due to the quadratic dependence of the X-ray emissivity on gas density. We thus expect little systematic bias in our measurement, in spite of the unrelaxed morphology of the system.

6.4 Baryon fraction

Fig. 11 shows the fraction of baryons in stars (green asterisks, see Section 6.2) and gas (red diamonds, see Section 6.3) as a function of cluster-centric radius. Both profiles exhibit trends typical for massive galaxy clusters, in which energy input from cluster mergers as well as feedback from active galactic nuclei (Nulsen et al. 2005) and galactic winds (Metzler & Evrard 1994) raise the entropy of the ICM. As a result, the hot gaseous atmosphere expands, and the gas fraction increases significantly with radius (e.g. Ettori & Fabian 1999; Vikhlinin et al. 2006). Conversely, the stars condense within the massive central galaxies, and the stellar fraction shows the opposite trend.

The total baryon fraction is shown by orange circles in Fig. 11. Within 1 Mpc from the cluster core, we measure a gas fraction of 0.072 ± 0.007 and a stellar fraction of 0.027 ± 0.004 . The total baryon fraction within this aperture is thus $f_{\text{bar}} = 0.099 \pm 0.008$. This value is 5σ below the cosmic baryon fraction measured by Planck Collaboration XVI (2014), $f_b = \Omega_b/\Omega_m = 0.1551 \pm 0.0055$ (black thick line in Fig. 11), and also discrepant with the cluster measurement of Mantz et al. (2014), $f_b = \Omega_b/\Omega_m = 0.14 \pm 0.02$, at more than 4σ confidence (dashed black thick line in Fig. 11).

This tension might be due to several factors. First, our analysis may have missed a significant fraction of the total stellar mass of the cluster, since the selection of cluster members considers mainly

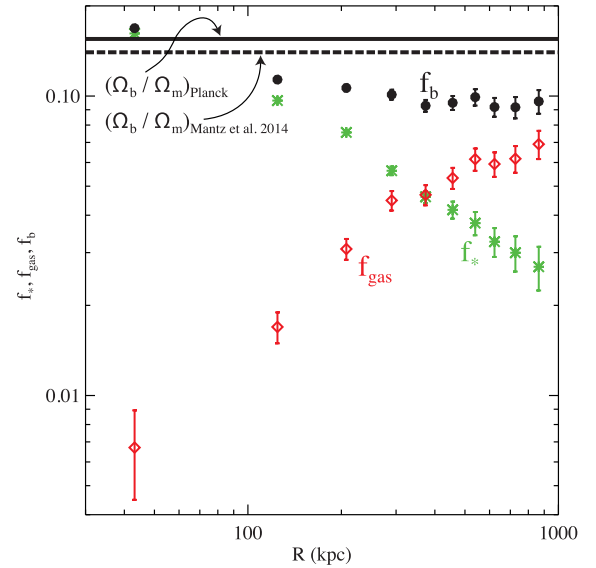


Figure 11. Fraction of stars, f_* , of gas, f_{gas} , and fraction of baryons, f_b present in MACSJ0416. We also plot the baryon fraction measured by Planck ($f_b = \Omega_b/\Omega_m = 0.1551 \pm 0.0055$), thick black line, and the one measured by Mantz et al. (2014) ($f_b = \Omega_b/\Omega_m = 0.14 \pm 0.02$), thick dashed black line.

red galaxies. The contribution from less massive (and fainter) star-forming cluster galaxies is not taken into account. In addition, intra-cluster light can account for 10–40 per cent of the total stellar mass (e.g. Gonzalez, Zaritsky & Zabludoff 2007; Giodini et al. 2009; Laganá et al. 2013). However, we estimate that, overall, the missing stellar content can contribute at most to 1 per cent of the total cluster mass. Another explanation is that a large fraction of the baryons resides outside the region sampled in our study. Recent studies (e.g. Simionescu et al. 2011; Eckert et al. 2013) have shown that the hot gas fraction continues to increase beyond R_{500} and eventually reaches the universal baryon fraction. It is likely that a significant fraction of the baryons indeed resides beyond 1 Mpc from the cluster centre, although the gas fraction of 7 per cent measured here is still significantly lower than the typical values measured in massive clusters around $R_{500} \sim 400$ kpc ($f_{\text{gas}} \sim 0.13$; Vikhlinin et al. 2006; Pratt et al. 2009). Finally, the total mass used to derive the baryon fraction could be overestimated by substructure along the line of sight. Our dynamical analysis of the member galaxies indeed revealed a difference of 800 km s^{-1} between C1 and C2, and the substructure S1 appears to be largely aligned with our line of sight (see Section 6.5). Since all of the aforementioned biases are known to be present but difficult to account for, we suggest that the deficit of baryons observed in Fig. 11 is probably due to a combination of these effects. In particular, it is likely that because of the merging activity a significant fraction of the baryons reside in the outskirts of the cluster, and given the presence of significant line-of-sight structure, the total lensing mass is also likely overestimated by some fraction. Becker & Kravtsov (2011) demonstrated that up to ~ 20 per cent scatter can be expected for WL mass measurements in the case of massive galaxy clusters, due to the presence of correlated and uncorrelated large-scale structures. The impact of these effects is however largely limited to large cluster-centric distances (> 3 Mpc), well beyond the value of ~ 1 Mpc from the cluster centre to which we map the mass distribution of MACSJ0416 in this study.

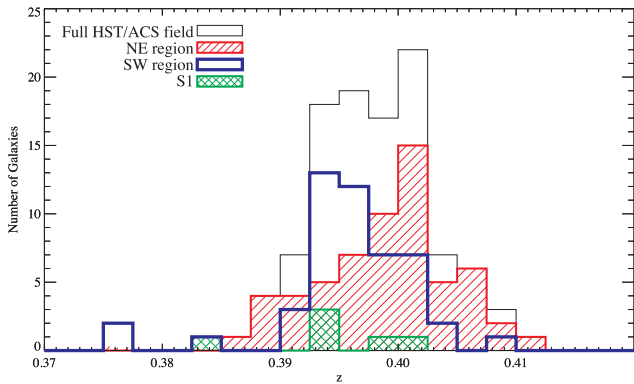


Figure 12. Distribution of spectroscopic galaxy redshifts in the field of MACSJ0416. The red and blue histograms show the different redshift distributions in the NE and SW sections of the ACS field, respectively, the green histogram shows the redshift distribution within the substructure S1 (see text for details).

6.5 Radial velocities of cluster galaxies

As reported in Section 6.1, the mass map derived by our joint SL and WL analysis reveals four significant mass concentrations: the two main merger components, as well as two smaller components (labelled S1 and S2 in Fig. 6) detected at more than 7σ confidence, that are located about 500 kpc (in projection) from the overall centre of the cluster.

As part of our attempt to clarify the nature and role of all four of these components in the assembly of MACSJ0416 (Section 7), we examine the redshift distribution of galaxies in the respective regions using the spectroscopic data described in Section 2.4. Using the *ROSTAT* package (Beers, Flynn & Gebhardt 1990), we find an overall redshift of 0.3980 for MACSJ0416 and a global velocity dispersion of 740 km s^{-1} , based on 106 spectroscopic redshifts. For a first global assessment, we divide the field of view along a boundary that runs perpendicular to the line connecting the main NE and SW cluster components (the apparent projected merger axis), intersecting it at its midpoint. Fig. 12 shows the overall distribution of spectroscopic galaxy redshifts (black), as well as, separately, the redshift distributions for the NE (red) and SW regions (blue) thus created. A two-sided Kolmogorov–Smirnov test yields a probability of less than 0.5 per cent for the hypothesis that the redshift distributions for the NE and SW regions are drawn from the same parent population. Their average redshifts are 0.3990 (NE) and 0.3966 (SW). The difference becomes more pronounced when the modes of the two redshift histograms are considered: 0.4013 (NE) and 0.3938 (SW).

Proceeding to the much less massive structures S1 and S2, we measure an average redshift of $z_{S1} = 0.3944$ from six cluster members with spectroscopic redshifts within 13 arcsec (~ 70 kpc), but are unable to estimate the redshift of S2 as we presently have no spectroscopic redshift for galaxies even within 15 arcsec of its centre.

7 THE MERGER GEOMETRY AND HISTORY OF MACSJ0416.1–2403

In the following, we combine our findings regarding the distribution of total gravitational mass, ICM, and cluster galaxies to derive a self-consistent picture of the evolutionary state and three-dimensional merger history of MACSJ0416.

7.1 Evidence

The bimodal distribution of galaxy redshifts shown in Fig. 12 and a projected separation of the two components of less than 300 kpc identify MACSJ0416 unambiguously as an active merger. The geometry of the collision, or even the answer to the question whether we observe MACSJ0416 before or after core passage, is not immediately obvious though.

Since MACSJ0416 has decoupled from the Hubble flow, the difference between the redshift distributions of galaxies in the NE and SW part of the cluster (Section 6.5) can safely be attributed to peculiar velocities. The difference between the means (modes) of the distributions corresponds to 500 (1600) km s^{-1} , while the difference in redshift between the BCGs of C1 and C2 (for simplicity we here adopt the nomenclatures of Table 2) implies a relative velocity of over 800 km s^{-1} – all remarkably high values compared to the global velocity dispersion of the system of 750 km s^{-1} . With peak collision velocities in massive mergers typically ranging from 1000 to 3000 km s^{-1} , the high relative radial velocity of the two subclusters strongly suggests a merger axis that falls markedly outside the plane of the sky. In a relative sense, C2 is thus moving towards us, while C1 is receding.

For an assessment of the direction of motion of the components of MACSJ0416 in the plane of the sky, as well as of the three-dimensional merger geometry in general, an inspection of the relative offsets (if any) between the collisional and non-collisional cluster components (the intracluster gas, and galaxies as well as dark matter) proves instructive. Binary head-on mergers (BHOM; Mann & Ebeling 2012) will feature a binary X-ray morphology before the collision, and a more unimodal morphology after (unless the merger axis falls very close to our line of sight). Offsets between gas and galaxies (and dark matter) will increase throughout the collision, as the non-collisional cluster components proceed unimpeded while the viscous ICM is shocked and slowed during the collision. Regardless of the merger axis, these offsets would be apparent in both participants in a BHOM. This is not the case for MACSJ0416. As shown in Section 6.3.1, collisional (gas) and non-collisional matter (galaxies and dark matter) coincide well for the NE component (C1 and X1), but are clearly displaced from each other (at 7σ significance) for the SW component (C2 and X2). Since non-collisional matter has to lead collisional matter in a merger, we infer that C2 is moving towards C1 (in projection).

An intriguing final piece of evidence is provided by the non-detection of structure S1 by *Chandra*. The lensing mass of S1 of $M_{S1}(R < 110 \text{ kpc}) = (4.22 \pm 0.56) \times 10^{13} h_{70}^{-1} M_{\odot}$ (Table 3) is substantial, and yet the mass derived for it from the upper limit to its X-ray luminosity of $L_X < 3.9 \times 10^{42} \text{ erg s}^{-1}$ is $M_{2500} < 10^{13} h_{70}^{-1} M_{\odot}$, based on the L_X – M relation for galaxy groups (Sun et al. 2009; Eckmiller, Hudson & Reiprich 2011). This apparent conflict, as well as the absence of hot gas in S1 as evinced by our tight upper limit to its gas fraction of $f_{\text{gas}} < 0.007$ (see Section 6.3.1), are easily explained though if S1 is in fact part of an unvirialized filamentary structure, almost aligned with our line of sight.

7.2 Tentative merger scenario

From the evidence compiled in the preceding section, we conclude that MACSJ0416 is (a) not a BHOM, i.e. the merger proceeds at significant impact parameter, rather than head-on, and (b) that the merger axis is greatly inclined with respect to the plane of the sky. MACSJ0416 is thus reminiscent of the merging system MACSJ0358.8–2955 for which Hsu et al. (2013) conclude that the

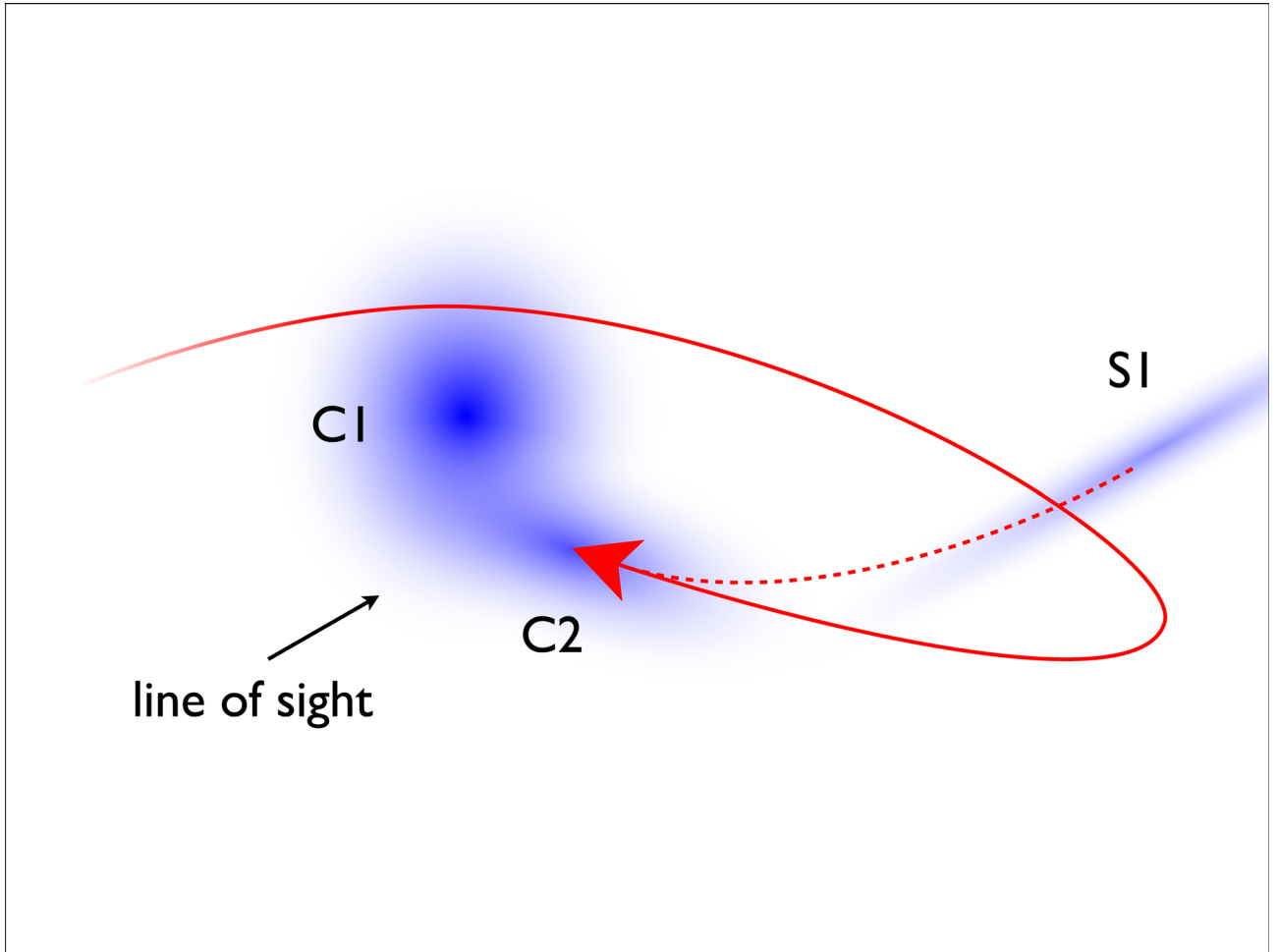


Figure 13. Schematic sketch of our hypothesized merger scenarios in a face-on view of the plane of the collision. The dotted red line marks the trajectory of C2 in our first scenario, in which the system is observed before its first core passage. The solid red line represents the second (preferred) scenario, in which C2 has passed C1 and is viewed after turn-around as it approaches the more massive cluster component for the second time. In this second scenario, filament S1 plays no significant role in the ongoing merger and may be positioned well before the cluster (along our line of sight), outside the area covered by this sketch.

lower mass component is likely moving along a trajectory that is curved towards our line of sight. The morphology of MACSJ0416 is very similar, except that the offset observed between the gas and the dark matter in the SW component (X2 and C2) goes in the opposite way as seen in MACSJ0358.8–2955.

We propose two alternative scenarios for the merger history of MACSJ0416, both of which are consistent with the present three-dimensional geometry as outlined above.³

Scenario #1: the SW component C2, observed near core passage, moves along a curved trajectory that originates in a large-scale filament, part of which is detected as substructure S1 by our lensing reconstruction of the mass distribution. Like for C2 itself, the mean redshift of galaxies near S1 of $z_{S1} = 0.3944$ is lower than that of the NE cluster C1, implying a radial velocity towards the observer of about 1000 km s^{-1} . The SW region of Fig. 6 contains a superposition of gas, galaxies, and dark matter from both this putative filament and C2, and hence the differences between the contours of gravitational mass, X-ray surface brightness, and cluster light reflect not only the

complex geometry and dynamical history of this system, but are also partly the result of the fact that lensing mass reconstructions collapse the mass of structures along the entire line of sight, virialized or not. Since, in this scenario, the merger of C1 and C2 resembles a ‘fly-by’ at the time of observation, the trajectory of the SW component has, so far, only grazed the core of the NE cluster which remains largely undisturbed, while ram pressure causes a disassociation between the ICM (X2) and the collisionless constituents (C2) of the approaching cluster (as seen in Fig. 6). A sketch of this scenario and of the trajectory of C2 is shown in Fig. 13 (dotted line) in a face-on view of the orbital plane.

Scenario #2: the SW component C2 approaches C1 just like in *Scenario #1*, but it does so for the second time. In this scenario, C2 does not originate in the large-scale filament S1; rather, C2 originally fell towards C1 from the opposite direction, passing the cluster core at a significant distance before turning around for its second approach. The core of C1 has been significantly disturbed by the earlier passage of C2, but any gas sloshing or shock heating induced by this interaction is almost imperceptible from our viewing angle, as the resulting cold fronts or shock fronts are not viewed edge-on, but are mainly projected on to the core of C1. The trajectory of C2 in this scenario is shown by the solid line in Fig. 13. In *Scenario #2*, filament S1 may be just behind the cluster (along our

³ Both of our scenarios are also in qualitative agreement with the one advanced by Diego et al. (2014) who propose that MACSJ0416 is merging along an axis that is only mildly inclined with respect to our line of sight.

line of sight), as indicated in Fig. 13, or far in front of MACSJ0416, well outside the virial regime. The X-ray evidence of multiphase gas in the core region of component C1 (Section 6.3.2) supports *Scenario #2*, which we thus presently favour.

More speculatively, and based largely on the apparent lack of a well formed X-ray core, we further propose that the SW component of MACSJ0416 is undergoing its own (minor) merger event. This second merger in the MACSJ0416 system proceeds again at high inclination with respect to the plane of the sky, is in the post-collision phase, and greatly disturbed the ICM, leading to the (unresolved) flat profile around X2. While this potential merger within the SW component of MACSJ0416 can be accommodated by either of our merger scenarios, it would fit more naturally into *Scenario #2* in which tidal forces during the first core passage may have aided the disruption of C2.

Although the currently available data do not allow us to clearly distinguish between these speculative scenarios, the tantalizing evidence of shock-heated gas in component C1 (Section 6.3.2) strongly supports *Scenario #2*. An opportunity to discriminate between *Scenarios #1* and *#2* will be provided by upcoming deep *Chandra*/ACIS-I observations of MACSJ0416. Ultimately reaching a cumulative exposure time of over 300 ks, these observations will be able to detect the presence of shocked gas near X1, as well as between X1 and C2 (clear sign of a previous interaction of the two components), as well as of cold fronts created by sloshing gas near X1. Dramatically improved photon statistics will yield ICM density and temperature maps also around X2 and stand to reveal the true dynamical history of this complex system.

8 SUMMARY AND CONCLUSION

We present a combined X-ray and optical analysis of the massive galaxy cluster MACSJ0416. Using the deep, high-resolution imaging data obtained by the HFF initiative, we model the mass distribution of the cluster over the *HST*/ACS field of view using a grid-based method that combines both SL and WL constraints. This lensing analysis is complemented by a study of the diffuse ICM, based on archival (16 ks + 37 ks) *Chandra* ACIS-I observation. Finally, we probe the distribution of mass along the line of sight using spectroscopic redshifts of 106 cluster members. We measure the following global properties: an overall galaxy velocity dispersion of 741 km s^{-1} and compelling evidence of bulk motions of $\sim 1000 \text{ km s}^{-1}$ along the line of sight; a total mass of $M(R < 950 \text{ kpc}) = (1.15 \pm 0.07) \times 10^{15} h_{70}^{-1} M_{\odot}$; an average mass-to-light ratio of $M_{*}/L_K = 0.99 \pm 0.03$; a gas mass of $M_{\text{gas}}(R < 1 \text{ Mpc}) = (8.6 \pm 0.7) \times 10^{13} h_{70}^{-1} M_{\odot}$; an ICM temperature of $kT = 11.0_{-1.3}^{+1.4} \text{ keV}$; an Fe abundance in the ICM of $0.20_{-0.08}^{+0.09} Z_{\odot}$; and a baryon fraction of $f_b(R < 1 \text{ Mpc}) = 0.099 \pm 0.008$ (5σ below the cosmological value estimated by the *Planck* mission). Importantly, our multiwavelength study also constrains the spatial distribution of dark and luminous matter in MACSJ0416 and reveals the presence of a massive ($M = (4.22 \pm 0.56) \times 10^{13} h_{70}^{-1} M_{\odot}$) but X-ray dark structure that we associate with a line-of-sight filament.

Using all observational evidence, we attempt to unravel the dynamical state and merger history of MACSJ0416. Central to our interpretation is the large offset in radial velocity between the two main cluster components, the fact that only the SW component shows a clear offset between collisional and non-collisional matter, and – possibly – our discovery of the aforementioned putative line-of-sight filament. We propose two alternative merger scenarios, the general geometry of which resembles that advanced by Hsu et al. (2013) for MACSJ0358.8–2955. Our two scenarios for

MACSJ0416 differ from one another primarily with regard to the pre- or post-collision state of the system. In either case, we assume that the NE component of MACSJ0416 is the more massive one. In our first merger scenario, the trajectory of the approaching SW component is slingshot-like, possibly originating in the aforementioned filament and passing the NE component at a significant impact parameter during the imminent first core passage. In our second scenario, the referred-to filament plays no role in the merger, which is much farther advanced than in our first scenario and in fact observed after turnaround of the SW component which now approaches the dominant NE component for the second time. The first passage significantly disturbed the core of the NE component, triggering gas sloshing of a modest cool core and shock heating of the gaseous ambient ICM. Since our spectral analysis of the ICM in the core region of the NE component finds tentative evidence of multiphase gas, we currently favour this second scenario. Travelling along a vector that is highly inclined with respect to the plane of the sky, the SW component appears internally disturbed and may be undergoing a minor merger of its own, regardless of the chosen scenario.

Our competing hypotheses regarding the merger history of MACSJ0416 make clear and testable predictions for the X-ray surface brightness and ICM temperature distribution around the NE component and for the region between the two subclusters. Forthcoming deep *Chandra* ACIS-I observations of MACSJ0416, capable of detecting the signature of sloshing gas and of temperature variations characteristic of ICM–ICM collisions thus hold great promise for a dramatically improved understanding of the formation history of this complex cluster lens.

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