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Floris van Der Tak, L. Chavarria, Fabrice Herpin, Friedrich Wyrowski,  
Malcolm Walmsley, Ewine van Dishoeck, The Wish Coordinating Team

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# Water in star-forming regions with *Herschel* (WISH)<sup>\*,\*\*</sup>

## IV. A survey of low-*J* H<sub>2</sub>O line profiles toward high-mass protostars

F. F. S. van der Tak<sup>1,2</sup>, L. Chavarría<sup>3,4,5</sup>, F. Herpin<sup>3,4</sup>, F. Wyrowski<sup>6</sup>, C. M. Walmsley<sup>7,8</sup>, E. F. van Dishoeck<sup>9,10</sup>, A. O. Benz<sup>11</sup>, E. A. Bergin<sup>12</sup>, P. Caselli<sup>13</sup>, M. R. Hogerheijde<sup>9</sup>, D. Johnstone<sup>14</sup>, L. E. Kristensen<sup>9,15</sup>, R. Liseau<sup>16</sup>, B. Nisini<sup>17</sup>, and M. Tafalla<sup>18</sup>

<sup>1</sup> SRON Netherlands Institute for Space Research, Landleven 12, 9747 AD Groningen, The Netherlands  
e-mail: vdtak@sron.nl

<sup>2</sup> Kapteyn Astronomical Institute, University of Groningen, The Netherlands

<sup>3</sup> Univ. Bordeaux, LAB, 33270 Floirac, France

<sup>4</sup> CNRS, LAB, 33271 Floirac, France

<sup>5</sup> Centro de Astrobiología (CSIC – INTA), 28850 Madrid, Spain

<sup>6</sup> Max-Planck-Institut für Radioastronomie, 53121 Bonn, Germany

<sup>7</sup> INAF – Osservatorio Astrofisico di Arcetri, 50125 Firenze, Italy

<sup>8</sup> Dublin Institute for Advanced Studies, Dublin 4, Ireland

<sup>9</sup> Leiden Observatory, Leiden University, The Netherlands

<sup>10</sup> Max-Planck-Institut für Extraterrestrische Physik, 85748 Garching, Germany

<sup>11</sup> Institute of Astronomy, ETH Zürich, Switzerland

<sup>12</sup> Department of Astronomy, University of Michigan, MI 48109, USA

<sup>13</sup> School of Physics and Astronomy, University of Leeds, LS2 9JT, UK

<sup>14</sup> Joint Astronomy Center, Hilo, Hawaii, HI 96720, USA

<sup>15</sup> Harvard-Smithsonian Center for Astrophysics, Cambridge, MA 02138, USA

<sup>16</sup> Onsala Space Observatory, Chalmers University of Technology, 41296 Göteborg, Sweden

<sup>17</sup> INAF – Osservatorio Astrofisico di Roma, Italy

<sup>18</sup> Observatorio Astronómico Nacional, 28049 Madrid, Spain

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### ABSTRACT

**Context.** Water is a key constituent of star-forming matter, but the origin of its line emission and absorption during high-mass star formation is not well understood.

**Aims.** We study the velocity profiles of low-excitation H<sub>2</sub>O lines toward 19 high-mass star-forming regions and search for trends with luminosity, mass, and evolutionary stage.

**Methods.** We decompose high-resolution *Herschel*-HIFI line spectra near 990, 1110 and 1670 GHz into three distinct physical components. Dense cores (protostellar envelopes) are usually seen as narrow absorptions in the H<sub>2</sub>O 1113 and 1669 GHz ground-state lines, the H<sub>2</sub>O 987 GHz excited-state line, and the H<sub>2</sub><sup>18</sup>O 1102 GHz ground-state line. In a few sources, the envelopes appear in emission in some or all studied lines, indicating higher temperatures or densities. Broader features due to outflows are usually seen in absorption in the H<sub>2</sub>O 1113 and 1669 GHz lines, in 987 GHz emission, and not seen in H<sub>2</sub><sup>18</sup>O, indicating a lower column density and a higher excitation temperature than the envelope component. A few outflows are detected in H<sub>2</sub><sup>18</sup>O, indicating higher column densities of shocked gas. In addition, the H<sub>2</sub>O 1113 and 1669 GHz spectra show narrow absorptions by foreground clouds along the line of sight. The lack of corresponding features in the 987 GHz and H<sub>2</sub><sup>18</sup>O lines indicates a low column density and a low excitation temperature for these clouds, although their derived H<sub>2</sub>O ortho/para ratios are close to 3.

**Results.** The intensity of the ground state lines of H<sub>2</sub>O at 1113 and 1669 GHz does not show significant trends with source luminosity, envelope mass, or evolutionary stage. In contrast, the flux in the excited-state 987 GHz line appears correlated with luminosity and the H<sub>2</sub><sup>18</sup>O line flux appears correlated with the envelope mass. Furthermore, appearance of the envelope in absorption in the 987 GHz and H<sub>2</sub><sup>18</sup>O lines seems to be a sign of an early evolutionary stage, as probed by the mid-infrared brightness and the  $L_{\text{bol}}/M_{\text{env}}$  ratio of the source.

**Conclusions.** The ground state transitions of H<sub>2</sub>O trace the outer parts of the envelopes, so that the effects of star formation are mostly noticeable in the outflow wings. These lines are heavily affected by absorption, so that line ratios of H<sub>2</sub>O involving the ground states must be treated with caution, especially if multiple clouds are superposed as in the extragalactic case. The isotopic H<sub>2</sub><sup>18</sup>O line appears to trace the mass of the protostellar envelope, indicating that the average H<sub>2</sub>O abundance in high-mass protostellar envelopes does not change much with time. The excited state line at 987 GHz increases in flux with luminosity and appears to be a good tracer of the mean weighted dust temperature of the source, which may explain why it is readily seen in distant galaxies.

**Key words.** stars: formation – ISM: molecules – astrochemistry

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\*\* Appendices are available in electronic form at <http://www.aanda.org>

## 1. Introduction

The water molecule is a key constituent of star-forming matter with a great influence on the formation of stars and planets<sup>1</sup>. In the gas phase, water acts as a coolant of collapsing dense interstellar clouds; in the solid state, it enhances the coagulation of dust grains in protoplanetary disks to make planetesimals; and as a liquid, it acts as a solvent bringing organic molecules together on planetary surfaces, which is a key step towards biogenic activity. The first role is especially important for high-mass star formation which depends on the balance between the collapse of a massive gas cloud and its fragmentation (Zinnecker & Yorke 2007). This balance depends strongly on the temperature (through the Jeans mass), and the sensitivity of the H<sub>2</sub>O abundance to the temperature, much larger than for CO, should make it a useful probe of the high-mass star formation process, which is mostly unexplored due to observational difficulties.

Interstellar H<sub>2</sub>O is well known from ground-based observations of maser emission at centimeter (22 GHz; Cheung et al. 1969), millimeter (183 GHz; Cernicharo et al. 1990), and submillimeter (325 GHz; Menten et al. 1990) wavelengths. The high intrinsic brightness of maser emission makes it useful as a signpost of dense gas and for kinematic studies of protostellar environments (e.g., Trinidad et al. 2003; Sanna et al. 2012a). With VLBI techniques, the proper motions and parallaxes of H<sub>2</sub>O masers can be measured to micro-arcsecond accuracy, leading to accurate distance estimates for star-forming regions as distant as ~10 kpc (Sanna et al. 2012b) and a revised picture of Galactic structure (Reid et al. 2009).

Thermal H<sub>2</sub>O lines are useful as probes of physical conditions and the chemical evolution of star-forming regions, but generally cannot be observed from the ground. Before *Herschel*, space-based submillimeter and far-infrared observations of H<sub>2</sub>O lines were made with ISO (Van Dishoeck & Helmich 1996), SWAS (Melnick & Bergin 2005) and *Odin* (Bjerkeli et al. 2009), but these data do not have sufficient angular resolution to determine the spatial distribution of H<sub>2</sub>O. In contrast, space-based mid-infrared and ground-based mm-wave observations of thermal H<sub>2</sub>O and H<sub>2</sub><sup>18</sup>O lines have high angular resolution but only probe the small fraction of the gas at high temperatures (Van der Tak et al. 2006; Watson et al. 2007; Jørgensen & van Dishoeck 2010; Wang et al. 2012).

The first results from the *Herschel* mission have demonstrated the potential of H<sub>2</sub>O observations of high-mass star-forming regions at high spatial and spectral resolution. Mapping of the DR21 region shows orders of magnitude variations in H<sub>2</sub>O abundances between various physical components (envelope, outflow, foreground clouds) due to freeze-out, evaporation, warm gas-phase chemistry, and photodissociation (Van der Tak et al. 2010). Multi-line observations of the high-mass protostar W3 IRS5 show broad emission as well as blueshifted absorption by H<sub>2</sub>O in the outflow, in particular in lines from excited states of H<sub>2</sub>O, underlining the importance of shock chemistry for H<sub>2</sub>O (Chavarría et al. 2010). Observations of the p-H<sub>2</sub>O and p-H<sub>2</sub><sup>18</sup>O ground state lines and the p-H<sub>2</sub>O first excited state line at 987 GHz toward four high-mass star-forming regions indicate H<sub>2</sub>O abundances ranging from  $5 \times 10^{-10}$  to  $4 \times 10^{-8}$  without a clear trend with the physical properties of the sources (Marseille et al. 2010b). Finally, observations of multiple H<sub>2</sub>O, H<sub>2</sub><sup>18</sup>O and H<sub>2</sub><sup>17</sup>O lines toward the massive star-forming region NGC 6334I

show high-excitation line emission from hot (~200 K) gas, and a H<sub>2</sub>O abundance ranging from  $\approx 10^{-8}$  in cold quiescent gas to  $4 \times 10^{-5}$  in warm outflow material (Emprechtinger et al. 2010). However, these papers study either single sources or small sets of sources, which makes a trend analysis inconclusive, if not impossible.

This paper presents observations of low-excitation lines of H<sub>2</sub>O and H<sub>2</sub><sup>18</sup>O toward 19 regions of high-mass star formation (Table 1). The velocity-resolved spectra show a mixture of emission and absorption due to various physical components, which we disentangle by fitting Gaussian profiles. The resulting line fluxes for the protostellar envelopes and the molecular outflows are compared with basic source parameters such as luminosity and mass, as well as to the evolutionary stage of the source as estimated in various ways.

To probe the bulk of the material in the protostellar environment, this paper focuses on the lowest excitation lines of H<sub>2</sub>O. The advantage of p-H<sub>2</sub>O is that its two lowest-excitation lines are close in frequency (Table 2) so that the telescope beam is similar in size. This coincidence allows us to make estimates of the column density and the excitation temperature of H<sub>2</sub>O which are almost independent of assumptions on the size or the shape of the object. In addition, we use observations of the o-H<sub>2</sub>O ground-state line at 1669 GHz to probe the kinematics of the sources, which have a high continuum brightness at this high frequency, so that absorption lines are easily detected. This paper does not discuss observations of the other o-H<sub>2</sub>O ground-state line at 557 GHz, because of the large difference in frequency and beam size. Maps of the 557 GHz line toward our sources will be presented elsewhere.

Table 1 presents the sources, which have been selected from several surveys (Molinari et al. 1996; Sridharan et al. 2002; Wood & Churchwell 1989; Van der Tak et al. 2000), where preference is given to nearby ( $\leq 2$  kpc) objects and “clean” objects where one source dominates the emission on  $\approx 30''$  scales, which is the relevant scale for our observations. Furthermore, the sources were selected to cover a range of evolutionary stages. In high-mass protostellar objects (HMPOs), the central star is surrounded by a massive envelope with a centrally peaked temperature and density distribution. These objects show signs of active star formation such as outflows and masers. This paper distinguishes mid-IR-bright and -quiet HMPOs, with a boundary of 10 Jy at 12  $\mu$ m, which may be an evolutionary difference (Van der Tak et al. 2000; Motte et al. 2007; López-Sepulcre et al. 2010). The presence of a “hot core” (submillimeter emission lines from complex organic molecules) or an “ultracompact HII region” (free-free continuum emission of more than a few mJy at a distance of a few kpc) are thought to be signposts of advanced stages of protostellar evolution. Hot cores occur when a star heats its surroundings such that ice mantles evaporate off dust grains which alters the chemical composition of the circumstellar gas. When the stellar atmosphere becomes hot enough for the production of significant ultraviolet radiation, the circumstellar gas is ionized and an ultracompact HII region is created.

This paper is organized as follows: Sect. 2 describes the observations and Sect. 3 their results. Section 4 presents a search for trends in the results with luminosity, mass or age, and Sect. 5 provides the discussion and our conclusions. In addition to this overview paper, our group is preparing detailed multi-line studies of individual sources.

## 2. Observations and data reduction

The sources were observed with the Heterodyne Instrument for the Far-Infrared (HIFI; De Graauw et al. 2010) onboard ESA’s

<sup>1</sup> This paper uses “water” to denote the chemical species, and H<sub>2</sub>O, H<sub>2</sub><sup>18</sup>O and HDO to denote specific isotopologues. Unless otherwise indicated, the gas phase is meant.

**Table 1.** Source sample.

Source	RA (J2000.0) h m s	Dec ° ' "	$V_{\text{LSR}}$ km s <sup>-1</sup>	$L_{\text{bol}}$ $L_{\odot}$	$d$ kpc	Ref.
Mid-IR-quiet HMPOs						
IRAS 05358+3543	05 39 13.1	+35 45 50	-17.6	$6.3 \times 10^3$	1.8	1, 1
IRAS 16272-4837	16 30 58.7	-48 43 55	-46.2	$2.4 \times 10^4$	3.4	2, 2
NGC 6334I(N)	17 20 55.2	-35 45 04	-4.5	$1.9 \times 10^3$	1.7	3, 4
W43 MM1	18 47 47.0	-01 54 28	+98.8	$2.3 \times 10^4$	5.5	5, 6
DR21(OH)	20 39 00.8	+42 22 48	-4.5	$1.3 \times 10^4$	1.5	7, 8
Mid-IR-bright HMPOs						
W3 IRS5	02 25 40.6	+62 05 51	-38.4	$1.7 \times 10^5$	2.0	9, 10
IRAS 18089-1732	18 11 51.5	-17 31 29	+33.8	$1.3 \times 10^4$	2.3	11, 12
W33A	18 14 39.1	-17 52 07	+37.5	$4.4 \times 10^4$	2.4	11, 30
IRAS 18151-1208	18 17 58.0	-12 07 27	+32.0	$2.0 \times 10^4$	2.9	13, 13
AFGL 2591	20 29 24.7	+40 11 19	-5.5	$2.2 \times 10^5$	3.3	14, 8
Hot molecular cores						
G327-0.6	15 53 08.8	-54 37 01	-45.0	$5.0 \times 10^4$	3.3	15, 16
NGC 6334I	17 20 53.3	-35 47 00	-7.7	$2.6 \times 10^5$	1.7	3, 4
G29.96-0.02	18 46 03.8	-02 39 22	+98.7	$3.5 \times 10^5$	6.0	17, 18
G31.41+0.31	18 47 34.3	-01 12 46	+98.8	$2.3 \times 10^5$	7.9	19, 20
Ultracompact HII regions						
G5.89-0.39 (W28A)	18 00 30.4	-24 04 02	+10.0	$5.1 \times 10^4$	1.3	21, 22
G10.47+0.03	18 08 38.2	-19 51 50	+67.0	$3.7 \times 10^5$	5.8	23, 20
G34.26+0.15	18 53 18.6	+01 14 58	+57.2	$3.2 \times 10^5$	3.3	24, 25
W51N-e1	19 23 43.8	+14 30 26	+59.5	$1.0 \times 10^5$	5.1	26, 27
NGC 7538-IRS1	23 13 45.3	+61 28 10	-57.4	$1.3 \times 10^5$	2.7	28, 29

**Notes.** The first reference is for the luminosity, the second for the distance. If the distance measurement is more recent than the luminosity estimate, the luminosity has been scaled to the new distance. The text uses “short” source names, which is the part preceding the + or - sign.

**References.** (1) Beuther et al. (2002); (2) Garay et al. (2002); (3) Sandell (2000); (4) Neckel (1978); (5) Nguyen Luong et al. (2011); (6) Lester et al. (1985); (7) Jakob et al. (2007); (8) Rygl et al. (2012); (9) Ladd et al. (1993); (10) Hachisuka et al. (2006); (11) Faúndez et al. (2004); (12) Xu et al. (2011); (13) Sridharan et al. (2002); (14) Van der Tak et al. (1999); (15) Urquhart et al. (2012); (16) Minier et al. (2009); (17) Cesaroni et al. (1998); (18) Pratap et al. (1999); (19) Mueller et al. (2002); (20) Churchwell et al. (1990); (21) Van der Tak et al. (2000) (22) Motogi et al. (2011); (23) Hunter et al. (2000); (24) Hatchell & van der Tak (2003) (25) Kuchar & Bania (1994); (26) Van Dishoeck et al. (2011); (27) Xu et al. (2009); (28) Sandell & Sievers (2004); (29) Moscadelli et al. (2009); (30) Immer et al. (2013).

**Table 2.** Observed lines.

Molecule	Transition	Frequency GHz	$E_{\text{up}}$ K	$\theta_{\text{mb}}$ arcsec	$\sigma_{\text{mb}}$ mK
H <sub>2</sub> O	1 <sub>11</sub> -0 <sub>00</sub>	1113.343	53.4	19.2	40
H <sub>2</sub> <sup>18</sup> O	1 <sub>11</sub> -0 <sub>00</sub>	1101.698	53.4	19.2	40
H <sub>2</sub> O	2 <sub>02</sub> -1 <sub>11</sub>	987.927	100.8	21.5	50
H <sub>2</sub> O	2 <sub>12</sub> -1 <sub>01</sub>	1669.905	114.4	12.7	75

**Notes.** The rms noise level in the last column is for HRS velocity resolution (0.5 MHz).

*Herschel* Space Observatory (Pilbratt et al. 2010) in the course of 2010 and 2011. Spectra were taken in double sideband mode using receiver bands 4b, 4a and 6b, respectively, for the 1<sub>11</sub>-0<sub>00</sub>, 2<sub>02</sub>-1<sub>11</sub> and 2<sub>12</sub>-1<sub>01</sub> lines (Table 2). The 1<sub>11</sub>-0<sub>00</sub> lines of H<sub>2</sub>O and H<sub>2</sub><sup>18</sup>O were observed simultaneously. All data were taken using the double beam switch observing mode with a throw of 2'5 to the SW, as part of the guaranteed time key program Water In Star-forming regions with *Herschel* (WISH; Van Dishoeck et al. 2011). See Appendix A for a detailed observing log.

Data were simultaneously taken with the acousto-optical Wide-Band Spectrometer (WBS) and the correlator-based High-Resolution Spectrometer (HRS), in both horizontal and vertical polarization. This paper focuses on the HRS data, which

cover 230 MHz bandwidth at 0.48 MHz ( $\sim 0.13$  km s<sup>-1</sup>) resolution, although for a few sources where the 1669 GHz line is too broad to fit in the HRS backend, the WBS data are used, which cover 1140 MHz bandwidth at 1.1 MHz ( $\sim 0.30$  km s<sup>-1</sup>) resolution. Spectroscopic parameters of the lines are taken from the JPL (Pickett et al. 1998) and CDMS (Müller et al. 2001) catalogs. The adopted FWHM beam sizes are taken from Roelfsema et al. (2012) and scaled to our observing frequencies. System temperatures are 340–360 K (DSB) at 1113 GHz, 400–450 K at 988 GHz, and 1400–1500 K at 1670 GHz. Integration times are 6.4 min at 1113 GHz (ON+OFF), 12 min at 988 GHz and 34 min at 1670 GHz.

The calibration of the data was performed in the *Herschel* Interactive Processing Environment (HIPE) version 6 or higher; further analysis was done within the CLASS<sup>2</sup> package, version of January 2011. The intensity scale was converted to  $T_{\text{mb}}$  scale using main beam efficiencies from Roelfsema et al. (2012), and linear baselines were subtracted. After inspection, data from the two polarizations were averaged together to obtain rms noise levels reported in the last column of Table 2. The calibration uncertainty is 15% at  $\sim 1000$  GHz and 20% at  $\sim 1700$  GHz (Ossenkopf & Shipman, priv. comm.)<sup>3</sup>.

<sup>2</sup> <http://www.iram.fr/IRAMFR/GILDAS>

<sup>3</sup> See <http://herschel.esac.esa.int/twiki/bin/view/Public/HifiCalibrationWeb?template=viewprint>

**Table 3.** SSB continuum flux densities (Jy/beam).

Source	987 GHz (304 $\mu\text{m}$ )	1108 GHz (271 $\mu\text{m}$ )	1669 GHz (180 $\mu\text{m}$ )
IRAS 05358	146	235	447
IRAS 16272	382	579	787
NGC 6334 I(N)	892	1100	887
W43 MM1	772	995	1073
DR21(OH)	896	1248	1678
W3 IRS5	546	760	1896
IRAS 18089	393	507	837
W33A	415	593	1216
IRAS 18151	106	217	308
AFGL 2591	310	467	830
G327-0.6	1206	1426	1621
NGC 6334I	2186	2601	4624
G29.96	510	688	1372
G31.41	492	684	948
G5.89	1020	1359	2908
G10.47	914	1340	2368
G34.26	1566	2272	3549
W51N-e1	1894	2388	2468
NGC 7538 IRS1	474	633	1037

**Notes.** The values for 1108 GHz are averages of measurements near the H<sub>2</sub>O 1113 GHz and the H<sub>2</sub><sup>18</sup>O lines.

### 3. Results

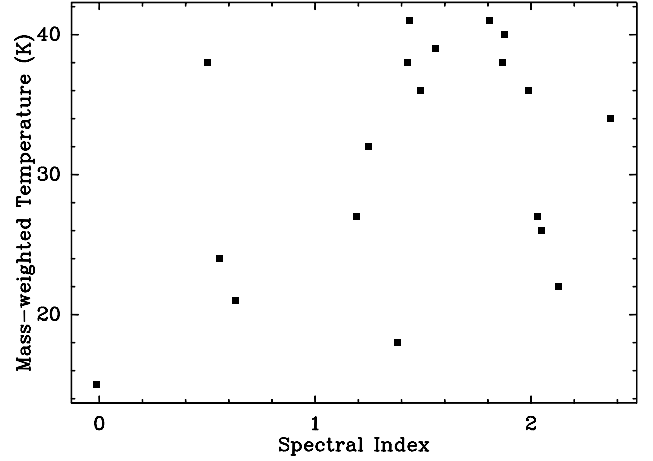
#### 3.1. Continuum emission

We have measured the continuum levels in the HIFI data by fitting zeroth- or first-order polynomials to the velocity ranges in the spectra without any appreciable line signal. In the cases of G327-0.6, NGC 6334I, and NGC 6334I(N), the HRS backend is not broad enough to contain the 1669 GHz line, so that we have used the WBS data instead. In general, the differences between continuum levels from the two backends (and within polarization channels from a given backend) are within 5% near 1000 GHz and within 10% at 1670 GHz, which is within the calibration uncertainty (Sect. 2). The flux densities  $S_\nu$  in Jy/beam were derived from the measured DSB continuum levels by

$$S_\nu = \frac{8kT'_A}{2\pi\eta_a D^2}$$

(Rohlfs & Wilson 2004, Eq. (7.18)) where the effective telescope diameter  $D = 3.28$  m and the aperture efficiency  $\eta_a \approx 0.65$  (Roelfsema et al. 2012) with a small frequency dependence due to Ruze losses. The factor 2 converts DSB to SSB signal, and  $T'_A = \eta_f T_A^*$  with the forward efficiency  $\eta_f = 0.96$  (Roelfsema et al. 2012).

As seen from Table 3, the continuum flux densities of our sources generally increase with frequency, as expected for thermal emission from warm dust. The free-free emission of our sources is  $\lesssim 1$  Jy which is negligible (Van der Tak & Menten 2005). For optically thick dust emission at a temperature  $T_d$ ,  $S_\nu = B_\nu(T_d)$  so that in the Rayleigh-Jeans limit, the spectral index  $\alpha$  (defined as  $S_\nu \propto \nu^\alpha$ ) equals 2, while for optically thin emission, the spectrum steepens to  $\alpha = 2 + \beta = 3-4$ , depending on the far-infrared dust opacity index  $\beta$ , which is 1–2 in most dust models (Ossenkopf & Henning 1994). However, at the high frequencies of our observations and the low average temperatures of our sources (15–40 K: Table 4), deviations from the Rayleigh-Jeans approximation are significant, which reduce these expected spectral indices by  $\approx 1$  at  $T_d = 40$  K and by  $\approx 2$  at  $T_d = 15$  K.



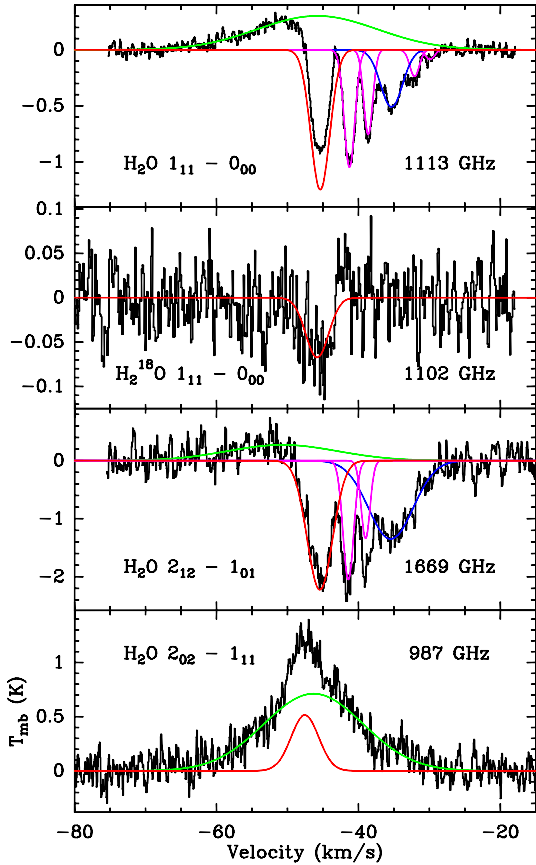
**Fig. 1.** Spectral indices of our sources, calculated from the HIFI continuum levels at 987 and 1669 GHz, versus the mass-weighted temperatures calculated from the models in Sect. 4.1.

**Table 4.** Parameters of continuum models.

Source	Input parameters		Output parameters	
	$R_{\text{out}}$ ( $10^4$ AU)	$p$	$M_{\text{env}}$ ( $M_\odot$ )	$\bar{T}^a$ (K)
IRAS 05358	5.4	1.5	142	22
IRAS 16272	17.0	1.5	2170	18
NGC 6334I(N)	7.2	1.3	3826	15
W43 MM1	14.7	1.5	7550	21
DR21(OH)	5.0	1.35	472	27
W3 IRS5	11.3	1.4	424	34
IRAS 18089	3.9	1.9	172	41
W33A	7.1	1.2	700	26
IRAS 18151	5.0	1.4	153	27
AFGL 2591	7.1	1.0	373	38
G327	8.0	1.2	2044	24
NGC 6334I	6.8	1.5	750	38
G29.96	9.6	1.4	768	40
G31.41	11.9	1.4	2968	32
G5.89	3.6	1.3	140	36
G10.47	6.0	1.4	1168	41
G34.26	8.1	1.4	1792	39
W51N-e1	14.0	1.4	4530	38
NGC 7538	6.4	1.0	433	36

**Notes.** <sup>(a)</sup> Mass-weighted temperature of the envelope model.

The spectral indices of our sources, calculated between 987 and 1670 GHz, are mostly between 1.2 and 2.4 (Fig. 1), which indicates that the emission is close to optically thick. The models in Sect. 4.1 also indicate continuum optical depths around unity at  $\sim 1000$  GHz. Four sources (NGC 6334I(N), W43 MM1, G327-0.6, and W51e) have lower spectral indices (between 0.0 and 0.5), which for compact sources would suggest low dust temperatures. For extended sources (size  $\gtrsim 15''$ ), the emission would be resolved at the highest frequencies, which would artificially flatten the spectrum. The model calculations in Sect. 4.1 indicate median source sizes of 17.6, 14.2, and 8.5'' FWHM at 987, 1113, and 1669 GHz, which is 83%, 75%, and 67% (i.e., most) of the corresponding beam size, suggesting that size effects play a minor role. The importance of dust temperature is also suggested by Fig. 1 which shows that the sources with low spectral indices have low mass-averaged temperatures as calculated from the models in Sect. 4.1. We conclude that our



**Fig. 2.** Gaussian decomposition of the observed H<sub>2</sub>O line profiles (after continuum subtraction) toward IRAS 16272. The envelope component is drawn in red, the broad outflow in green, the narrow outflow in blue, and foreground clouds in purple.

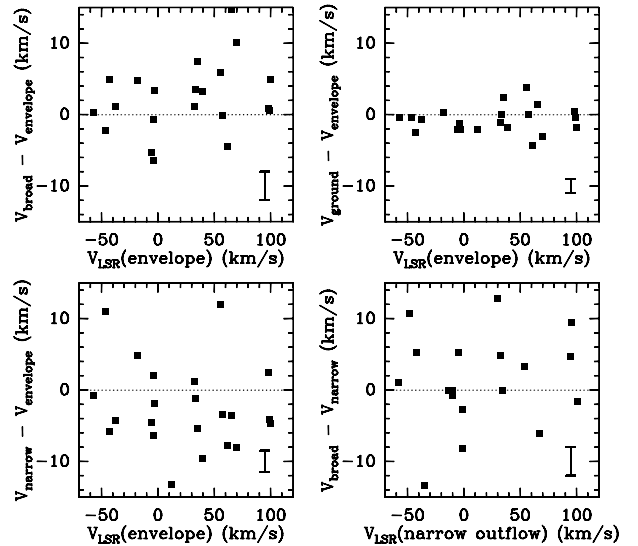
observed variations in the spectral index of the HIFI continuum emission are not due to variations in the source size, but rather to a low dust temperature, both directly and through deviation from the Rayleigh-Jeans law.

### 3.2. H<sub>2</sub>O lines

Figures B.1–B.19 in Appendix B show the observed line profiles. All four lines are detected toward all nineteen sources (except H<sub>2</sub><sup>18</sup>O toward IRAS 18151), and the line profiles are a mixture of emission and absorption features. Analogous to the case of DR21 (Van der Tak et al. 2010), three physical components appear from these observations (Fig. 2).

First are the protostellar envelopes (also sometimes called dense cores) which usually appear in all four lines in absorption at an LSR velocity known from ground-based millimeter-wave emission line data (Fig. 3, top right); the widths are also similar to the ground-based values. Several sources show two envelope components, due to binarity (e.g., W3 IRS5) or infall/outflow profiles (e.g., W33A, G34.26). The general appearance in absorption indicates a high H<sub>2</sub>O column density and a low excitation temperature for this component, which is further discussed in Sect. 3.4.

Second are the molecular outflows which usually appear in 1113 and 1669 GHz absorption and 987 GHz emission at velocity offsets of 5–10 km s<sup>-1</sup> from the envelope signal and with widths of 10–20 km s<sup>-1</sup> FWHM or sometimes more, up to 40 km s<sup>-1</sup>. Most sources show two outflow components: a



**Fig. 3.** Central velocities of the envelope, narrow outflow, and broad outflow components of the observed H<sub>2</sub>O lines compared with each other and with the systemic velocity from ground-based data. In the bottom right-hand corner, the typical error bar on the velocities is shown.

broad ( $\Delta V \sim 20$  km s<sup>-1</sup>) and a narrower ( $\Delta V \sim 10$  km s<sup>-1</sup>) feature. In some cases, the width of the narrow outflow feature is similar to that of the corresponding envelope, but the comparison to the ground-based velocity and the presence or absence of H<sub>2</sub><sup>18</sup>O emission or absorption allows for an unambiguous assignment. The lack of H<sub>2</sub><sup>18</sup>O signal and the appearance in 987 GHz emission suggest that the outflows have lower H<sub>2</sub>O column densities and higher excitation temperatures than the envelopes. Section 3.5 discusses the outflow component in detail.

Third are foreground clouds, which appear as narrow ( $\Delta V < 5$  km s<sup>-1</sup>) 1113 and 1669 GHz absorption features at velocities offset by 5–50 km s<sup>-1</sup> from the envelope signal. The lack of signal in the 987 GHz and H<sub>2</sub><sup>18</sup>O lines indicates a low H<sub>2</sub>O column density and a low excitation temperature for these components, without exception. The number of foreground clouds within the HRS bandwidth ranges from zero toward IRAS 05358 and NGC 7538 to seven toward W51e. Many of these foreground clouds are known from ground-based mm-wave observations of ground-state molecular lines such as CS, HCO<sup>+</sup> and HCN (Greaves & Williams 1994; Godard et al. 2010) and from [HI] 21 cm observations (Fish et al. 2003; Pandian et al. 2008). Other foreground clouds have first been seen with HIFI, such as the  $V = +24$  km s<sup>-1</sup> cloud toward W51e which is seen in HF (Sonnentrucker et al. 2010) but not in our H<sub>2</sub>O data, and the  $V = +13$  km s<sup>-1</sup> cloud toward AFGL 2591 which is seen in H<sub>2</sub>O and HF but not in CO (Emprechtinger et al. 2012; Van der Wiel et al. 2013; Choi et al., in prep.). Some foreground clouds are intervening objects on the line of sight, while others represent the “outer envelopes” of large-scale molecular cloud complexes, such as the  $V = +65$  km s<sup>-1</sup> cloud toward W51e (Koo 1997) and two clouds toward NGC 6334 (Van der Wiel et al. 2010). See Sect. 3.6 for further discussion of the foreground clouds.

Our decomposition of the line profiles is similar to the description of H<sub>2</sub>O line profiles toward low-mass protostars in terms of “narrow”, “medium” and “broad” components (Kristensen et al. 2010), where narrow = envelope, medium = narrow outflow, and broad = broad outflow. Because of the large range in luminosity and mass within our sample, we base the assignment of features to envelope or outflow not only on the line

width, but also on the appearance in emission or absorption in the various lines.

### 3.3. Profile decomposition and flux extraction

We have measured the fluxes ( $\int T_{\text{mb}} dV$ ; in emission) and absorbances ( $\int \tau dV$ ; in absorption) of the lines by fitting Gaussians to the observed profiles for each line independently. Figure 2 shows IRAS 16272 as an example of the procedure, and the results are presented in Table A.2 for the envelopes, Table A.3 for the outflows, and Table A.4 for the foreground clouds. Our fitting procedure is similar to that by Flagey et al. (2013), where the reader can find further details.

Since some of the emission and absorption features in our data have non-Gaussian shapes, we have verified the results of the Gaussian fitting procedure by a “moment” analysis, which just adds up the flux in all spectrometer channels within a given velocity range. This alternative method extracts the flux of a spectral component independent of its shape, but it cannot handle blending of components, which occurs frequently in our data, unless further assumptions are made. Another disadvantage of the moment method is that the definitions of the velocity ranges are often somewhat arbitrary. After some experimenting, we conclude that fitting Gaussians gives more reliable results than taking moments, if we accept that in some cases, one physical component (usually the outflow) is represented by several (up to three) Gaussians. Another consequence of the non-Gaussian line shape is that the extracted central position and the width of a given component vary somewhat from line to line: up to  $2 \text{ km s}^{-1}$  for the envelopes and up to  $5 \text{ km s}^{-1}$  for the broad outflow component. The uncertainties on the line positions and widths in Tables A.2, A.3 and A.4 represent this line-to-line variation, whereas the uncertainties on the fluxes and absorbances are dominated by the overall calibration uncertainty (Sect. 2), which is always larger than the formal error of the Gaussian fit caused by the deviation of the observed profile from a Gaussian shape.

For absorption features, Tables A.2–A.4 give the absorbance (=the velocity-integrated optical depth):

$$\int \tau dV = -1.06 \Delta V \ln \left( \frac{T_c - T_1}{T_c} \right)$$

where  $T_1$  is the peak depth of the feature,  $T_c$  is the SSB continuum temperature, and  $\Delta V$  is the FWHM line width. For small optical depth, the error on  $T_1$  propagates linearly into the absorbance, but if the absorption is almost saturated ( $T_1 \rightarrow T_c$ ), the error on the absorbance is much larger than that on the measured absorption depth. Cases where the error exceeds the absorbance due to saturation, which occurs mostly for envelope components, are marked with lower limits in Tables A.2–A.4.

In the frequent case that absorption and emission features are partially blended, we have estimated the line contribution to the absorption background from the spectra, and added the result to the continuum brightness. This procedure assumes that the absorber is located in front of both the continuum and the line emitting source, which is plausible since in most cases, the absorption depth significantly exceeds the continuum brightness alone.

The model calculations in Sect. 4.1 indicate median source sizes of 17.6, 14.2, and 8.5 arcsec FWHM at 987, 1113, and 1669 GHz, which is 83%, 75%, and 67% of the corresponding beam size. The filling factor would be the square of this fraction

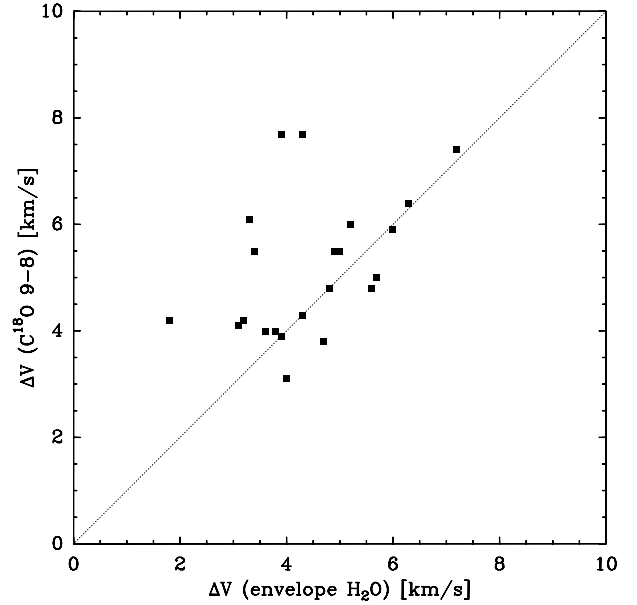


Fig. 4. Line widths (FWHM) of the envelope components of the observed  $\text{H}_2\text{O}$  lines compared to the widths of the  $\text{C}^{18}\text{O}$  9–8 line measured with HIFI (San José García et al. 2013). Error bars are smaller than the plotting symbols.

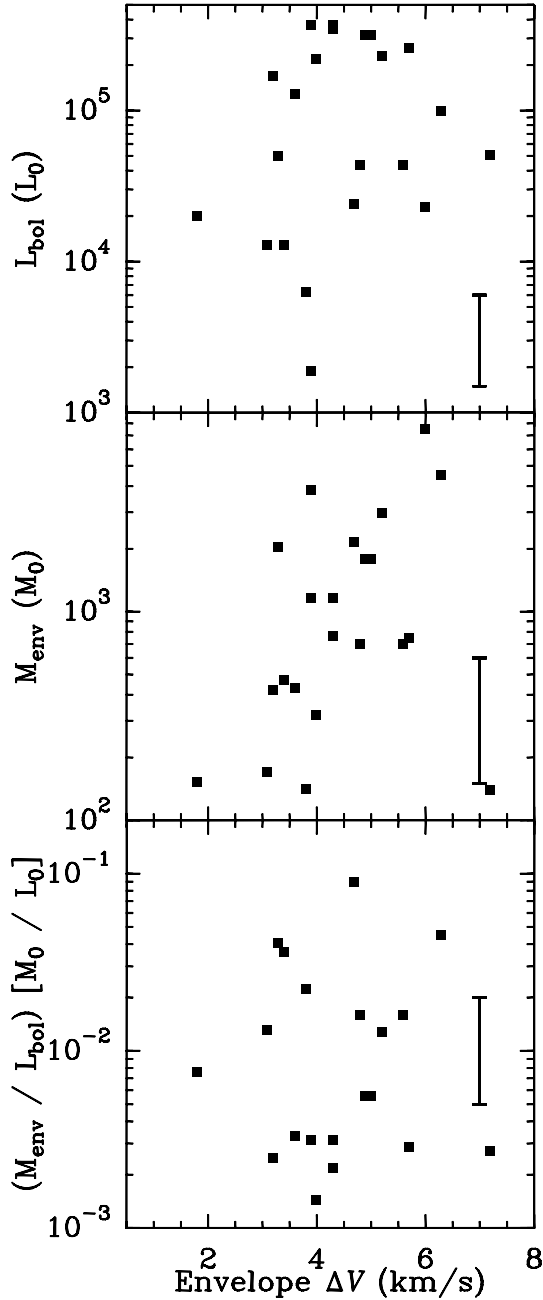
which is 68%, 56%, and 45% respectively, so that our absorption line measurements will be affected by covering factor effects if the absorbers are much smaller in size than the beam. In this case, the beam-averaged optical depth would be below the source-averaged value.

### 3.4. Envelopes

The  $\text{H}_2\text{O}$  line spectra of all sources show envelope components, and the velocities and widths of these components in Table A.2 agree well with previous measurements for these sources in other molecules from the ground (Fig. 3). The widths also mostly agree with the widths of the  $\text{C}^{18}\text{O}$  9–8 line observed with HIFI in a similar beam (Fig. 4); the few cases where  $\Delta V(\text{C}^{18}\text{O}) > \Delta V(\text{H}_2\text{O})$  may be influenced by an outflow contribution to the  $\text{C}^{18}\text{O}$  9–8 line. The measured  $\text{H}_2\text{O}$  line widths range over a factor of 4, and show a clear correlation with the envelope masses calculated in Sect. 4.1 (Fig. 5).

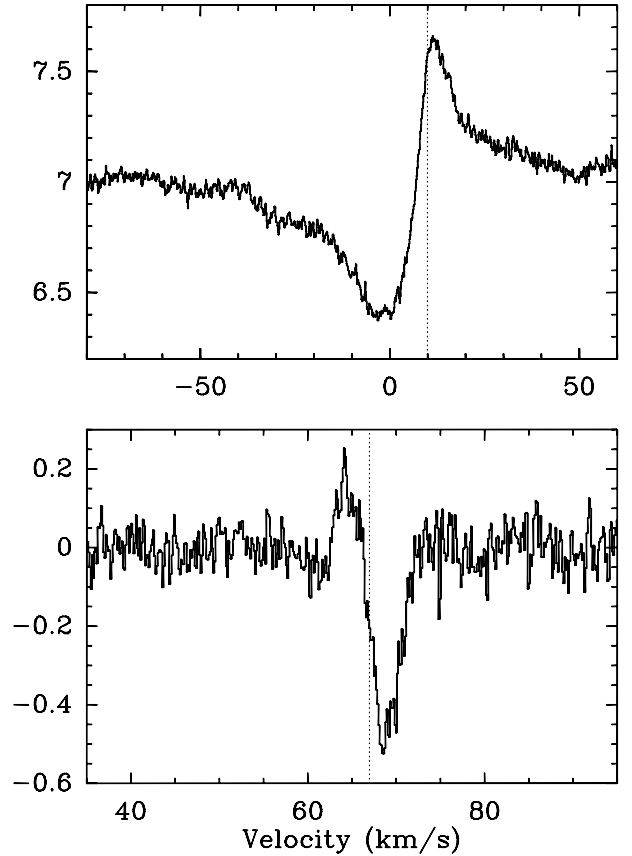
The appearance of the envelopes in the four lines differs markedly from source to source, which can be grouped into four types. Seven envelopes appear in absorption in all four lines: G31.41, DR21OH, G327, NGC 6334I(N), IRAS 05358, W43MM1, and one of the G34.26 envelopes. Six envelopes show emission in the 987 GHz line and absorption in the other lines: G10.47, IRAS 18089, W51e, NGC 6334I, IRAS 16272, and one of the W33A envelopes. A further six envelopes appear in  $\text{H}_2\text{O}$  987 GHz and in  $\text{H}_2^{18}\text{O}$  emission, and in  $\text{H}_2\text{O}$  1113 and 1669 GHz absorption: G29.96, NGC 7538, W3 IRS5, G5.89, and the second envelopes of W33A and G34.26. Finally, one source, AFGL 2591, appears in emission in all four lines. One curious case is IRAS 18151, where the non-detection of the  $\text{H}_2^{18}\text{O}$  line leaves ambiguous whether it belongs in the second or third of the above source types; or perhaps it is in a transition from the second to the third type.

One possible origin for the diverse appearance of the envelope features is related to temperature, as appearance in emission (absorption) requires an excitation temperature above (below)



**Fig. 5.** Envelope line width measured from our spectra versus bolometric luminosity (*top*), envelope mass (*middle*), and mass/luminosity ratio (*bottom*). Sources with two envelope features are shown with two data points. In the bottom right-hand corner, the typical uncertainty in the masses, luminosities, and mass/luminosity ratios is shown.

the continuum brightness. In this scenario, pure absorption sources would have the most massive and coldest envelopes, sources where the 987 GHz line appears in emission would be somewhat warmer, sources where both H<sub>2</sub>O 987 GHz and H<sub>2</sub><sup>18</sup>O appear in emission would be even warmer, and the pure emission source would be the warmest. Alternatively, the appearance of the envelope features in emission or absorption may be a matter of continuum brightness, as detectability in absorption requires a background source whereas emission does not. In particular, the presence of outflow cavities in the HIFI beam would decrease the continuum compared with a beam filled with envelope emission. However, the continuum flux densities in Table 3 do not support this option: the four envelope types above have



**Fig. 6.** *Top:* H<sub>2</sub><sup>18</sup>O line profile toward G5.89, showing a P Cygni profile. *Bottom:* H<sub>2</sub><sup>18</sup>O line profile toward G10.47, showing an inverse P Cygni profile. The bottom spectrum is continuum subtracted. Vertical scale is  $T_{\text{mb}}$  in K, and the dotted line denotes the systemic velocity from ground-based data.

very similar median continuum brightness levels. Outflow cavities may also affect the H<sub>2</sub>O line signals, but the net effect may either be emission or absorption, depending on the difference between the H<sub>2</sub>O excitation temperature and the dust temperature, and on the relative location of dust and water along the line of sight. We conclude that heating effects are more likely than geometric effects to be responsible for the variation in envelope appearance between our sources.

Some sources show two envelope features in some or all H<sub>2</sub>O line profiles, which in certain cases may be due to the presence of common-envelope binaries at  $\sim 1000$  AU separation, as suggested for W3 IRS5 (Chavarría et al. 2010). In other cases, the lines have P Cygni or inverse P Cygni profiles, which are due to expansion and infall, respectively (Fig. 6). Such profiles have also been observed in the H<sub>2</sub>O line spectra of low-mass protostars (Kristensen et al. 2012) and for the low-mass pre-stellar core L1544 (Caselli et al. 2012). Interestingly, the (inverse) P Cygni signature is usually not seen in all lines toward a given source. The sources G10.47 and G34.26 show inverse P Cygni profiles in the H<sub>2</sub>O 987 GHz and H<sub>2</sub><sup>18</sup>O lines, suggesting that infall motion only takes place in their inner envelopes. For G34.26, this result is confirmed by SOFIA observations of red-shifted NH<sub>3</sub> line absorption at 1810 GHz (Wyrowski et al. 2012), and for G10.47 by SMA observations of vibrationally excited HCN (Rolfs et al. 2011). In contrast, G5.89 and DR21OH show P Cygni profiles in the H<sub>2</sub>O 1113 and 1669 GHz lines and in H<sub>2</sub><sup>18</sup>O, but not in the H<sub>2</sub>O 987 GHz line, suggesting that expanding motions only take place in their outer envelopes. In W33A,



where only the  $\text{H}_2^{18}\text{O}$  line may show a P Cygni profile, the expanding motion may take place further inside the envelope. For W43 MM1, the low- $J$  lines do not show infall or outflow motions, consistent with the observation by Herpin et al. (2012) that only high- $J$  lines show infall motions in this source, i.e., that infall is confined to the inner envelope. Similar qualitative kinematic variations with radius have been observed in HCN lines toward the Sgr B2 envelope, where the emission peak shifts from blue to red with increasing  $J$ , suggesting infall in the outer and expansion in the inner envelope (Rolfs et al. 2010).

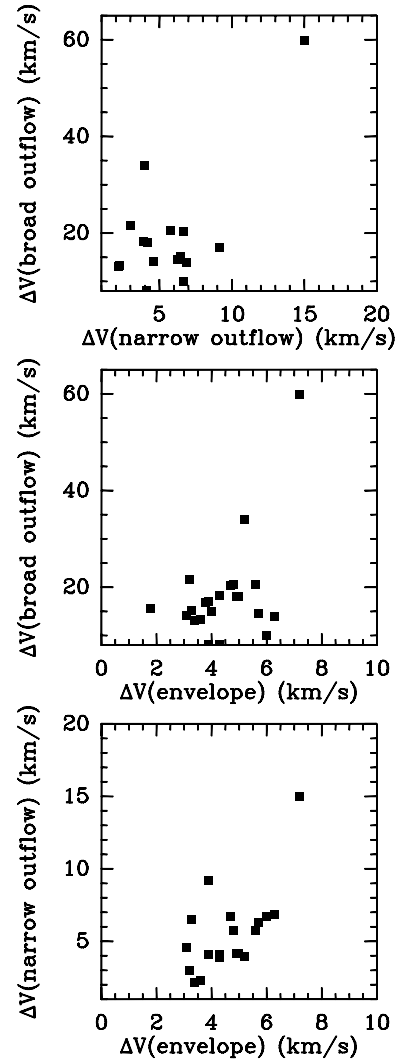
### 3.5. Outflows

The outflows usually show up in the  $\text{H}_2\text{O}$  line profiles as two components: a broad ( $\Delta V \sim 10\text{--}20 \text{ km s}^{-1}$ ) emission component visible in all lines, and a narrower ( $\Delta V \sim 5\text{--}10 \text{ km s}^{-1}$ ) absorption component seen in all but the 987 GHz line. Exceptions are AFGL 2591, IRAS 18151 and IRAS 05358, which only show a broad component, and G5.89, which has an exceptionally broad and complex line profile; this source is known to have a very massive and powerful outflow (Acord et al. 1997). For the sources DR21OH, W3 IRS5 and NGC 7538, the widths of the narrower components are so small ( $2\text{--}3 \text{ km s}^{-1}$ ) that these may be “second envelopes”, as also seen toward W33A and G34.26, and as also advocated by Chavarría et al. (2010) for W3 IRS5. The mm-wave continuum emission from these sources indeed shows multiple spatial components (Woody et al. 1989; Van der Tak et al. 2005), sometimes with mid-infrared counterparts (Megeath et al. 2005; De Wit et al. 2009).

For most of our sources, the broad outflow component appears in emission in all four lines, but there are some exceptions. For AFGL 2591, NGC 6334I, W51e and G34.26, the broad outflow component appears in absorption in the 1113 and 1669 GHz lines, and in emission in the 987 GHz line; for G10.47, this component appears in emission in the 1113 and 987 GHz lines, and in absorption in the 1669 GHz line. Toward W51e, a “very broad” component appears in emission in all four lines between  $+30$  and  $+90 \text{ km s}^{-1}$ . It is unclear in which category G31.41 belongs, since the broad outflow component is only detected in the 987 GHz line, perhaps because in the ground state lines, emission and absorption cancel each other out. The shape of the broad outflow component on the 987 GHz line profile is often non-Gaussian, but the outflow contribution can be modeled with two Gaussians of about the same height and width. In this case the velocity in Table A.3 is the mean of the values for the two Gaussians, while the width and the flux are the sums.

The narrow outflow component usually appears in absorption in the  $\text{H}_2\text{O}$  1113 and 1669 GHz lines and the  $\text{H}_2^{18}\text{O}$  line, and is not seen in the  $\text{H}_2\text{O}$  987 GHz line, but again, exceptions occur. For W3 IRS5, G327 and G31.41, the narrow outflow component appears in 1113 and 1669 GHz absorption and in 987 GHz emission; for IRAS 18089 and G34.26, this component appears in 1113 and 1669 GHz absorption and in  $\text{H}_2^{18}\text{O}$  emission, while for G10.47, it appears in emission in all four lines. Unlike the envelopes, the outflow absorptions are usually not saturated, with G327 as only exception.

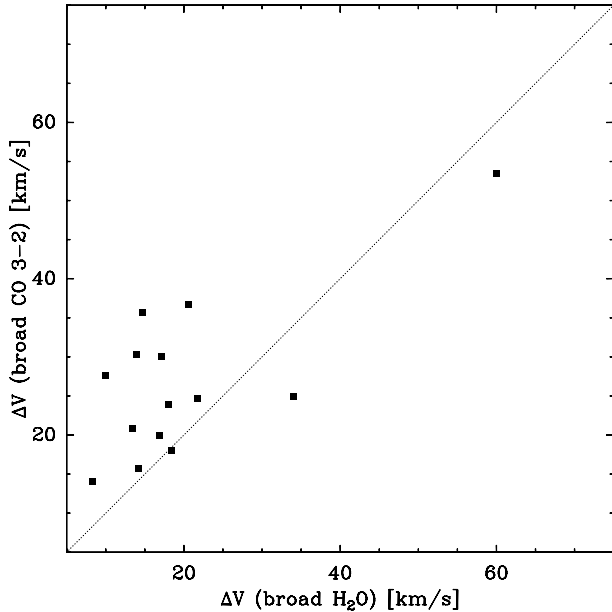
We have used the measured velocities of the outflow components to test if our sources have a preferred orientation in the sky. In particular, for a face-on geometry, the outflow would be directed toward us, and thus appear blueshifted relative to the envelope component if seen in absorption. The narrow outflow component is the one that is usually seen in absorption, and Fig. 3 shows that this component is indeed usually blueshifted relative to the envelope. The exceptions are



**Fig. 7.** Line widths (FWHM) of the envelope, narrow outflow, and broad outflow components of the observed  $\text{H}_2\text{O}$  lines compared with each other. Error bars are smaller than the plotting symbols.

IRAS 16272, DR21(OH), G29.96, and W51e, and Fig. 2 illustrates the case of IRAS 16272, especially for the 1669 GHz line. The widths of these redshifted absorption features are too large to be due to infall motions, and we suggest that these sources have binaries or other deviations from centrosymmetric geometry that allow their receding outflow lobes to show in absorption. In addition, Fig. 7 shows that there are no clear trends between the widths of the envelope and outflow components.

Figure 8 compares the widths of the broad outflow components to the widths observed in CO 3–2 with the JCMT (San José García et al. 2013). While sources with broader  $\text{H}_2\text{O}$  outflows also seem to have broader CO outflows, the width for any given source is larger in CO than in  $\text{H}_2\text{O}$  by about  $10 \text{ km s}^{-1}$ . This result is contrary to the case of low-mass protostars, where the  $\text{H}_2\text{O}$  line widths exceed those of CO (Kristensen et al. 2012). For our sources, the  $\text{H}_2\text{O}$  may be confined to denser gas where velocities are lower. Alternatively, the difference is excitation: our  $\text{H}_2\text{O}$  values are averages over ground-state and excited-state lines, while the outflow line widths increase with  $J$  or  $E_{\text{up}}$ . Indeed, the widths of the excited-state  $\text{H}_2\text{O}$  lines are very close to those measured in CO, as also found by Chavarría et al. (2010) and Herpin et al. (2012).



**Fig. 8.** Line widths (FWHM) of the broad outflow components of the observed H<sub>2</sub>O lines compared to the widths of the CO 3–2 lines measured with the JCMT (San José García et al. 2013). Error bars are smaller than the plotting symbols.

### 3.6. Foreground clouds

The foreground clouds appear as narrow ( $\Delta V = 1\text{--}4\text{ km s}^{-1}$ ) absorption features in the 1113 and 1669 GHz line spectra of our sources. While one cloud also shows H<sub>2</sub><sup>18</sup>O absorption, the general lack of signal in the H<sub>2</sub>O 987 GHz and H<sub>2</sub><sup>18</sup>O lines indicates a low column density and a low volume density for these clouds. The last two columns of Table A.4 list estimates of the H<sub>2</sub>O column densities and ortho-para ratios of the foreground clouds. These values were obtained from the velocity-integrated optical depths in the table by assuming that in these clouds, all H<sub>2</sub>O molecules are in their ortho and para ground states, i.e., that excitation is negligible. As discussed by Emprechtinger et al. (2013), this assumption generally holds for column density estimates from the 1113 and 1669 GHz lines, but fails for the H<sub>2</sub>O 557 GHz line.

The H<sub>2</sub>O column densities of the foreground clouds are found to lie mostly in the range  $1 \times 10^{12}\text{--}1 \times 10^{13}\text{ cm}^{-2}$ , with a few outliers at the upper and lower end. These values are comparable to earlier estimates (Van der Tak et al. 2010; Lis et al. 2010) and indicate that the H<sub>2</sub>O abundance in these clouds is limited by photodissociation. The derived ortho/para ratios for H<sub>2</sub>O lie mostly in the range between 3 and 6, again with a few outliers to higher and lower values. Our derived o/p ratios are very similar to the values by Lis et al. (2010) and Herpin et al. (2012). Comparison to the detailed multi-line study by Flagey et al. (2013) for the 5 overlapping sources suggests that our estimated column densities and o/p ratios for individual clouds have uncertainties of a factor of  $\sim 2$ , which is an upper limit since not exactly the same positions were observed.

The ortho-para ratio of H<sub>2</sub>O is expected to rise from  $\approx 1.0$  at low temperatures ( $\approx 15\text{ K}$ ) to  $\approx 3$  at high temperatures ( $\geq 50\text{ K}$ ) as shown by Mumma et al. (1987). Even though the values derived here are probably uncertain to a factor of  $\sim 2$ , our results are inconsistent with the low-temperature limit. If the H<sub>2</sub>O molecules were formed on cold (10 K) grain surfaces, their spin temperature does not reflect these conditions, and may instead be set by

the grain temperature shortly before the H<sub>2</sub>O ice mantle evaporated (at  $T \approx 100\text{ K}$ ). Alternatively, the H<sub>2</sub>O molecules may have formed in the warm ( $T_{\text{kin}} \approx 50\text{ K}$ ) gas phase by ion-molecule reactions, or the o/p ratio of the evaporated H<sub>2</sub>O may have equilibrated with that of H<sub>2</sub> in the warm gas phase in reactions with H<sub>3</sub><sup>+</sup>.

## 4. Analysis and discussion

### 4.1. Mass estimation

Estimates of the envelope masses exist in the literature for many of our sources. However, these mass estimations were performed by different authors, using different techniques and based on data from different telescopes. In order to have a homogeneous mass determination over our sample, we derive the envelope mass by comparing models of the continuum spectral energy distribution (SED) with the available observations for each source.

We use a modified version of the 3D continuum radiative transfer code HOCHUNK3D (Whitney et al. 2013; Robitaille 2011, hereafter WR) to find the envelope temperature profile and derive a model SED for each source. Our modified version of the code uses dust opacities from Ossenkopf & Henning (1994), model 5 with thin ice mantles and a gas density of  $10^6\text{ cm}^{-3}$ , suitable for high-mass protostars (Van der Tak et al. 1999), and allows to vary the radial density power-law exponent  $p$ . The WR code allows for a detailed description of the protostellar environment in terms of a star, a disk, a cavity and an envelope. To decrease the number of free parameters, we use a simple spherically symmetric model with no cavity and no disk. This simplification has its main effect in the near to mid-infrared region of the spectrum, while the sub-millimeter and millimeter-wave regions, where the dust is mainly emitting, will stay almost unchanged.

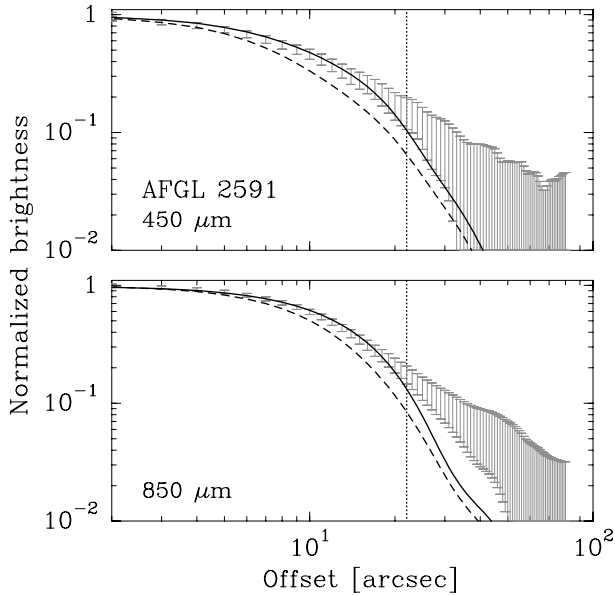
The envelope size and density exponent  $p$  are derived from archival sub-millimeter images from the JCMT/SCUBA<sup>4</sup> and APEX/LABOCA (Siringo et al. 2009) instruments. For the envelope size, we choose the radius of the contour at the  $3\sigma$  noise level, while for the density exponent, we use the WR code to create synthetic sub-millimeter maps for various  $p$ -values. We determine the best-fit density exponent by minimizing the difference between the observed and synthetic brightness profiles (see Fig. 9). The synthetic maps are created using the corresponding instrumental pass-band and convolved with the instrumental beam profile including the error-beam.

With the outer radius and the density profile fixed, the only remaining parameters are the envelope mass and the bolometric luminosity. The observed fluxes from the literature correspond to the peak emission at the source position. In most cases, the beam size of the observed fluxes is smaller than the source size. To correct for the missing flux, we normalize the observed fluxes by the ratio between the flux from the whole model and the flux from a region as big as the beam size. As an example, Fig. 10 shows the observed and derived SED for the source IRAS 16272 as well as the temperature profile. Table 4 lists the adopted and the derived physical parameters for each source, and Appendix C presents the results in graphical form.

### 4.2. Correlations with luminosity and mass

To understand the origin of the observed H<sub>2</sub>O lines, we now investigate if the measured line fluxes show trends with the

<sup>4</sup> The James Clerk Maxwell Telescope (JMCT) is operated by the Joint Astronomy Centre on behalf of the Science and Technology Facilities Council of the United Kingdom, the Netherlands Organisation for Scientific Research, and the National Research Council of Canada.

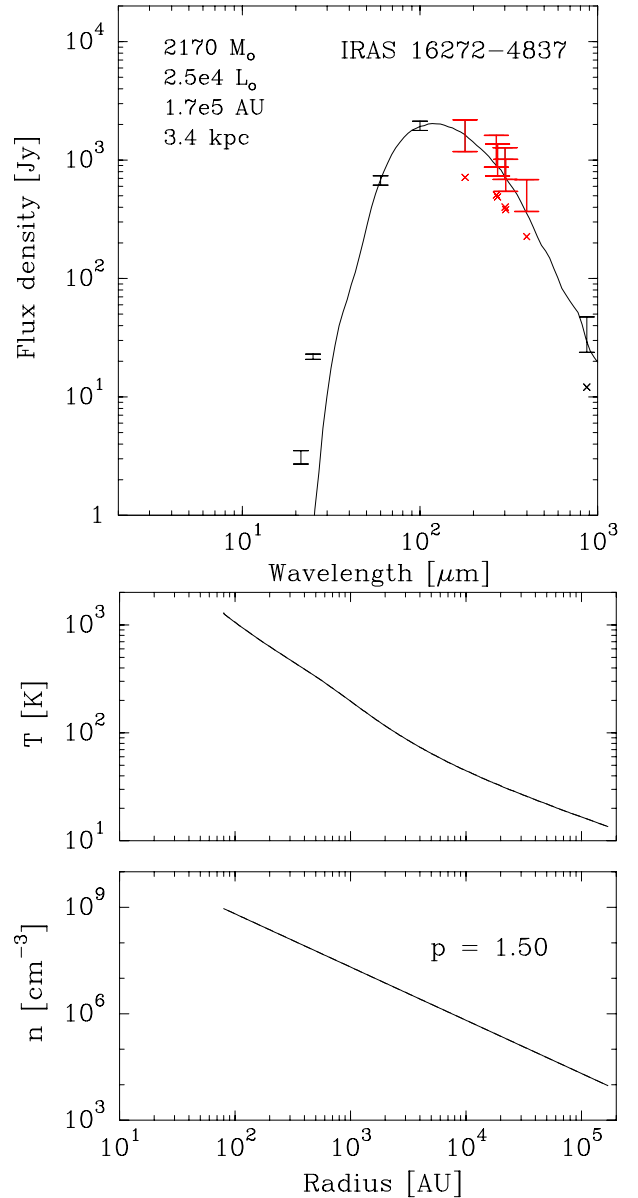


**Fig. 9.** Submillimeter brightness profiles of AFGL 2591 measured with JCMT/SCUBA (error bars), and synthetic brightness profiles derived from WR models and convolved by the SCUBA beams for  $p = 1.0$  (solid line) and  $p = 1.5$  (dashed line). The vertical dotted line indicates the adopted source radius. For this source, we choose a density profile of  $p = 1.0$ .

luminosity of the source  $L_{\text{bol}}$  (Table 1) and the mass of the envelope  $M_{\text{env}}$  (Table 4). In addition, we have searched for trends with distance but did not find any. For the luminosity, we use the results of the model fit in Sect. 4.1 so that the source flux at all wavelengths refers to the same area on the sky. Figures D.1–D.3 in Appendix D show scatter plots of the line fluxes of the envelope and the narrow and broad outflow components versus  $L_{\text{bol}}$  and  $M_{\text{env}}$ . In these figures, absorption is plotted as negative emission, and “double envelope” features are plotted as two data points. We identify possible trends by visual inspection of these plots, and test these (if applicable) by computing the correlation coefficient  $r$  and the probability of false correlation  $P$ , which depends on  $r$  and the number of data points  $N$  (see Appendix of Marseille et al. 2010a).

For the envelope component, the visual inspection reveals two trends. First, the 987 GHz line flux increases with  $L_{\text{bol}}$ , but only if the line appears in emission. This trend is not a correlation, but more a “threshold effect”: strong 987 GHz line emission is only seen at high luminosities ( $L_{\text{bol}} \gtrsim 10^5 L_{\odot}$ ). Second, the  $\text{H}_2^{18}\text{O}$  line flux decreases with increasing  $M_{\text{env}}$ , but again not as a correlation but in the sense that the line is seen in emission at low  $M_{\text{env}}$  and absorption for high  $M_{\text{env}}$ . Such behavior is expected if the line optical depth, and thus the  $\text{H}_2\text{O}$  column density, increases with envelope mass. An optical depth of  $\sim$ unity for the  $\text{H}_2^{18}\text{O}$  line is confirmed by detections of  $\text{H}_2^{17}\text{O}$  line emission in several of our sources (Herpin et al. 2012; Choi et al., in prep.). Apparently the average  $\text{H}_2\text{O}$  abundance is rather constant across all envelopes in our source sample. The studies mentioned in Sect. 1 indicate outer envelope abundances of  $\sim 10^{-9}$ – $10^{-8}$ .

The increase of the 987 GHz line flux with  $L_{\text{bol}}$  is probably due to a higher temperature of the inner envelopes of sources with a higher luminosity. This idea is confirmed by the trend between 987 GHz line flux with far-infrared continuum flux density (Fig. 11). Such a trend is seen both at 60 and at 100  $\mu\text{m}$ , and like the trend with  $L_{\text{bol}}$ , is more a threshold effect than a correlation. The large beam size of IRAS may influence these

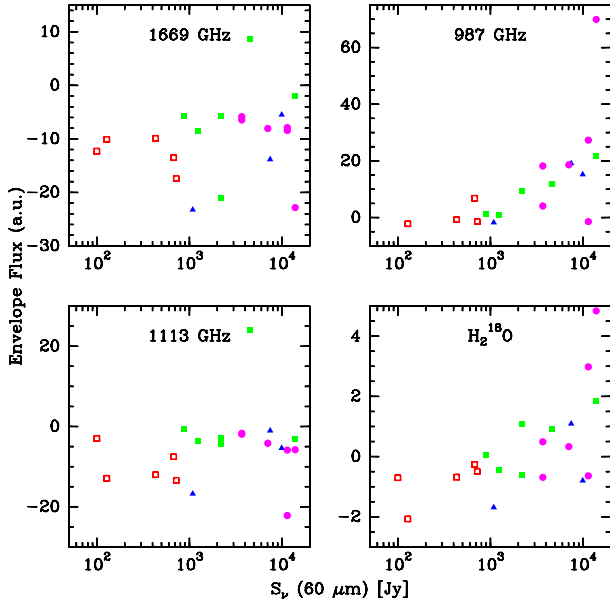


**Fig. 10.** Top: model (continuous line) and observed SED fluxes (error bars) for IRAS 16272. The SED points are normalized by the source size for  $\lambda > 100$  microns, as shown by the “X” symbols. Continuum flux densities from our HIFI data are labeled in red. Bottom: temperature and density profiles for IRAS 16272.

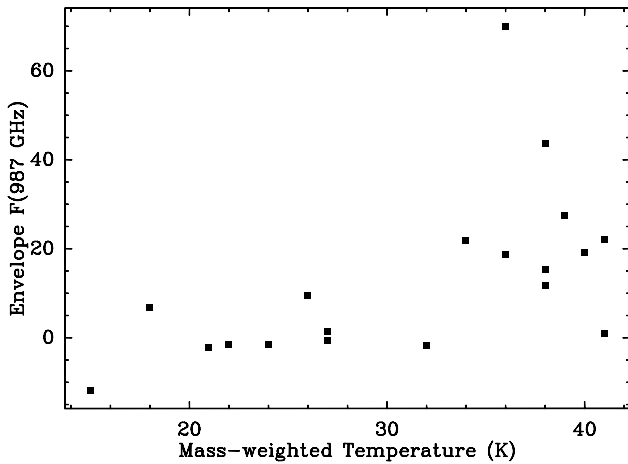
trends quantitatively, but probably has no strong qualitative effect because our sources were selected to be relatively isolated (Sect. 1). Far-infrared pumping could influence the results for individual  $\text{H}_2\text{O}$  lines, but is unlikely to dominate here, because the effect is seen in several lines.

Further evidence that the  $\text{H}_2\text{O}$  987 GHz line flux is due to envelope heating is presented in Fig. 12, which is a plot of the flux in this line versus the mass-weighted temperature from the envelope models in Sect. 4.1. While a 1:1 correlation cannot be claimed, the points show a clear trend toward higher line fluxes at higher average temperatures.

Figures D.2 and D.3 show a similar analysis for the narrow and broad outflow components in the  $\text{H}_2\text{O}$  line profiles of our sources. The most significant trend is that the  $\text{H}_2^{18}\text{O}$  line flux of the broad outflow component decreases with increasing  $L_{\text{bol}}$  ( $r = -0.683$ ,  $P = 3\%$ ), but only in the 11 cases where this component



**Fig. 11.** Envelope line flux versus  $60\ \mu\text{m}$  flux density. Uncertainties are  $\approx 10\%$  for HIFI data and  $\approx 20\%$  for IRAS data. Open red and filled green squares are mid-IR-quiet and -bright protostellar objects, filled blue triangles are hot molecular cores, and filled purple circles are ultracompact HII regions.

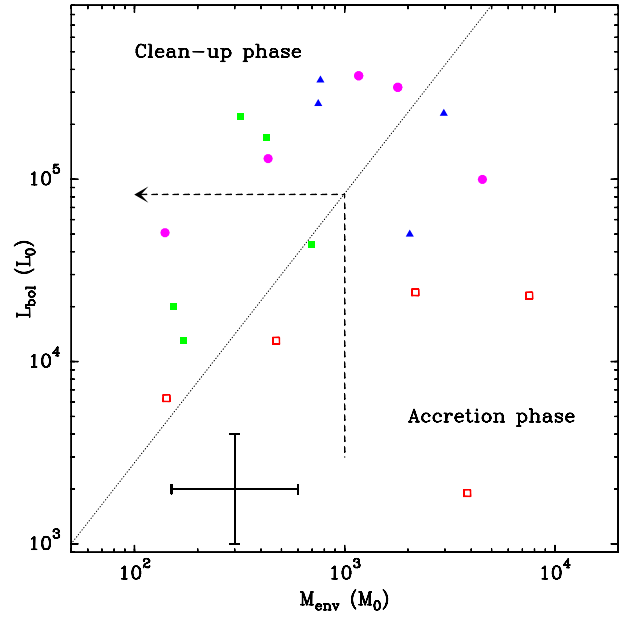


**Fig. 12.** Envelope  $\text{H}_2\text{O}$  987 GHz line flux versus mass-weighted temperature in our envelope models.

is detected in  $\text{H}_2^{18}\text{O}$ . Second, the 987 GHz line flux of the broad outflow component appears to increase with increasing  $M_{\text{env}}$ , but with  $r = 0.174$  and  $P = 14\%$ , this trend is less significant than the first one. The line fluxes of the narrow outflow component do not show any clear trends with  $L_{\text{bol}}$  or  $M_{\text{env}}$ .

The lack of evolutionary trends in the  $\text{H}_2\text{O}$  line fluxes of our sources is in contrast with the study of López-Sepulcre et al. (2011), who observed the SiO  $J = 2-1$  and  $3-2$  lines toward 57 high-mass star-forming regions, and found that the SiO line luminosity drops with evolutionary stage as measured by increasing values of  $L_{\text{bol}}/M_{\text{env}}$ . This drop suggests a decline of the jet activity with the evolutionary state of the source, as also observed for low-mass sources (Bontemps et al. 1996), although excitation effects also may play a role.

One reason for this difference may be that even the  $\text{H}_2\text{O}$  line signals from the outflows may be somewhat optically thick, so that saturation masks any evolutionary trends. Therefore, we



**Fig. 13.** Bolometric luminosity versus envelope mass for our sources, where the dotted line divides the two phases of high-mass protostellar evolution suggested by Molinari et al. (2008) and the dashed arrow is a typical evolutionary track. The typical uncertainty of a factor of 2 in both luminosity and mass is indicated at the bottom. Symbols are as in Fig. 11.

consider next trends of the line width with  $L_{\text{bol}}$  and  $M_{\text{env}}$ . The line width is less sensitive to the optical depth than the line intensity, and has the additional advantage that the appearance of the line in emission or absorption does not influence its measurement. Section 3.4 already shows the trend of the envelope  $\Delta V$  with  $M_{\text{env}}$ , which with  $r = 0.407$  and  $P = 4\%$  for  $N = 22$  is likely real.

Figures D.4 and D.5 show scatter plots of the widths of the narrow and broad outflow components versus  $L_{\text{bol}}$  and  $M_{\text{env}}$ . The only significant trend that appears is that  $\Delta V$  of the narrow outflow component increases with  $M_{\text{env}}$ . A relation of this component with the envelope has been found before for low-mass protostars (Kristensen et al. 2012), and suggests that the narrow outflow component is due to outflow-envelope interactions.

### 4.3. Evolutionary trends

We now discuss several proposed evolutionary indicators for high-mass star formation and test them on our source sample. We will use the ratio of  $L_{\text{bol}}$  over  $M_{\text{env}}$  as primary evolutionary indicator, because of previous success (Van der Tak et al. 2000; Motte et al. 2007; López-Sepulcre et al. 2010), because it is a continuous variable, and because it has a physical basis (declining accretion activity and envelope dispersal). In particular, Molinari et al. (2008) have proposed an evolutionary sequence for high-mass protostars in terms of two parameters: the envelope mass and the bolometric luminosity. They suggest a two-step evolutionary sequence: protostars first accrete mass from their envelopes, and later disperse their envelopes by winds and/or radiation. In the diagram in Fig. 13, their calculated evolutionary tracks start at the bottom, rise upward until they reach the dotted line, and then proceed horizontally to the left. The positions of our sources in this diagram suggest that the subsample of mid-IR-quiet protostellar objects is the youngest, and that the other subsamples contain more evolved sources. The same

effect is seen in Figs. D.1–D.3, where the mid-IR-quiet protostellar objects occupy the left (low- $L_{\text{bol}}$ ) sides of the diagrams. On the other hand, the presence of a hot molecular core and/or an ultracompact HII region is unrelated to mid-infrared brightness and  $L_{\text{bol}}/M_{\text{env}}$  ratio, at least for our sample. The occurrence of these phenomena may depend more on absolute luminosity than on evolutionary stage.

To test the use of H<sub>2</sub>O line profiles as evolutionary indicators, we use the appearance of the envelope component in the H<sub>2</sub>O ground-state line profiles, as introduced in Sect. 3.4. In Type 1 sources, all lines appear in absorption; Type 2 sources show emission in the 987 GHz line and absorption in the others; in Type 3 sources, both the 987 GHz and the H<sub>2</sub><sup>18</sup>O lines appear in emission; and Type 4 sources show all lines in emission. While the Type 1 sources all seem to be in early evolutionary stages as probed by a low  $L_{\text{bol}}/M_{\text{env}}$  ratio, the other source types are all mixed. Thus, an H<sub>2</sub>O line profile with pure absorption appears to be a sign of an early evolutionary stage, whereas emission is a sign of a more evolved stage.

The appearance of the broad and narrow outflow components in the H<sub>2</sub>O ground-state line profiles differs between our sources, and for each component, four types may be defined as for the envelopes (Sect. 3.5). While these emission/absorption types may be useful to trace protostellar evolution, our data do not allow to test this hypothesis, since most of our sources belong to just one of these types, with the broad outflow appearing in absorption and the narrow outflow in emission. This issue may be revisited after HIFI data become available for a larger sample of sources outside the WISH program.

Besides the shape of the line profile, the fluxes and widths of the H<sub>2</sub>O line may also be evolutionary indicators. The bottom panels of Figs. D.1–D.3 show scatter plots of our observed line fluxes in the envelope and outflow components versus the ratio  $M_{\text{env}}/L_{\text{bol}}$  which as discussed above is sometimes used as evolutionary indicator for high-mass protostars. No trends appear, and the same is true for the plots of the line width versus  $M_{\text{env}}/L_{\text{bol}}$  (Figs. 5, D.4 and D.5).

## 5. Summary and conclusions

This paper presents spectroscopic observations of low- $J$  H<sub>2</sub>O transitions observed toward 19 regions of high-mass star formation. The line profiles show contributions from protostellar envelopes, bipolar outflows, and foreground clouds. These contributions, both in emission and in absorption, are disentangled by fitting multiple Gaussians to the line profiles.

Our observations show that the ground state lines of H<sub>2</sub>O are powerful tracers of the various physical and kinematic components of the high-mass protostellar environment, as well as of foreground clouds. The envelope components of the ground state lines of H<sub>2</sub>O at 1113 and 1669 GHz do not show significant trends with bolometric luminosity, envelope mass, or far-infrared flux density or color, indicating that these lines trace the outer parts of the envelopes where the effects of star formation are hardly noticeable. On the other hand, broad wings due to outflows are clearly seen in these lines, which are a clear effect of star formation. The competing effects of line emission and absorption in the ground-state lines lead to low integrated line fluxes, which may explain the difficulty in detecting these lines in distant galaxies.

Unlike the ground-state lines, the flux in the excited-state 987 GHz line increases systematically with luminosity and with the far-infrared flux density, both at 60 and at 100  $\mu\text{m}$ . This line therefore appears to be a good tracer of the average temperature

of the source. Indeed, this line and other excited-state lines have proven to be more easily detected at high redshift (Omont et al. 2011; Van der Werf et al. 2011).

Furthermore, the H<sub>2</sub><sup>18</sup>O ground-state line turns from emission into absorption with increasing envelope mass. This trend indicates that the H<sub>2</sub>O abundances in protostellar envelopes do not change much with time.

The appearance of the H<sub>2</sub>O line profiles in our sample changes with evolutionary state, in the sense that the youngest sources (as probed by mid-infrared brightness and  $L_{\text{bol}}/M_{\text{env}}$  ratio) show pure absorption profiles for all lines, while for more evolved sources, one or more of our studied lines appears in emission. This effect is visible for the envelope component but not for the outflows. In addition, the presence of infall or expansion, as probed by (inverse) P Cygni profiles, does not appear related to evolutionary stage.

In the future, our team will make detailed estimates of the H<sub>2</sub>O abundances in these sources as a function of radius, based on HIFI observations of  $\approx 20$  lines of H<sub>2</sub>O, H<sub>2</sub><sup>18</sup>O and H<sub>2</sub><sup>17</sup>O. The *Herschel*/PACS instrument will be used to study higher-excitation H<sub>2</sub>O lines and to improve estimates of the bolometric temperatures of our sources, which may act as evolutionary indicator. Maps of selected lines with HIFI and PACS will clarify the spatial distribution of H<sub>2</sub>O, both on the scale of protostellar envelopes and outflows as on larger (protostellar cluster) scales. On smaller scales, ALMA will provide images of high-excitation H<sub>2</sub>O and H<sub>2</sub><sup>18</sup>O lines and thus zoom into the close vicinity of the central protostars.

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## Appendix A: Observation log and line flux tables

**Table A.1.** List of observations.

Source	H <sub>2</sub> O & H <sub>2</sub> <sup>18</sup> O	H <sub>2</sub> O	H <sub>2</sub> O
	1 <sub>11</sub> -0 <sub>00</sub>	2 <sub>02</sub> -1 <sub>11</sub>	2 <sub>12</sub> -1 <sub>01</sub>
IRAS 05358	206 126	204 510	203 954
IRAS 16272	214 419	203 168	192 584
NGC 6334I(N)	206 385	204 520	214 454
W43-MM1	191 670	191 616	192 575
DR21(OH)	194 794	195 026	192 569
W3 IRS5	191 658	191 612	192 566
IRAS 18089	229 883	215 911	216 696
W33A	191 638	191 636	192 576
IRAS 18151	229 880	218 211	218 909
AFGL 2591	197 973	195 019	192 570
G327	214 425	203 170	192 585
NGC 6334I	206 383	204 519	214 455
G29.96	191 668	191 617	192 574
G31.41	191 671	191 615	192 573
G5.89	229 888	218 120	216 694
G10.47	229 884	215 914	216 695
G34.26	191 674	194 995	192 572
W51e	194 801	195 014	192 571
NGC 7538	191 663	201 599	200 758

**Notes.** The preceding 1342 has been omitted from all obsIDs.

**Table A.2.** Central velocities, FWHM widths, and integrated line fluxes measured for the protostellar envelopes.

Source	$V_{\text{LSR}}$	$\Delta V$	1113 GHz	H <sub>2</sub> <sup>18</sup> O	1669 GHz	987 GHz
IRAS 05358	-17.9(6)	3.8(14)	-13.4	-0.49	>17.4	-1.39
IRAS 16272	-45.9(7)	4.7(20)	-7.51	-0.26	>13.5	+6.83
NGC 6334I(N)	-2.4(5)	3.9(4)	-2.96	-0.70	>12.3	-11.8
W43 MM1	+100.5(7)	6.0(15)	-12.9	-2.07	-10.1	-2.17
DR21(OH)	-3.3(4)	3.4(13)	>12.0	-0.68	>9.9	-0.69
W3 IRS5	-37.7(12)	3.2(6)	-3.06	+1.84	-2.00	+21.9
IRAS 18089	+33.7(11)	3.1(11)	-3.47	-0.45	-8.44	+0.90
W33A	+35.1(12)	5.6(13)	-4.34	-0.61	>21.0	<0.60
<sup>a</sup>	+39.3(6)	4.8(14)	-2.85	+1.09	-5.66	+9.46
IRAS 18151	+33.1(2)	1.8(10)	-0.52	<0.07	>5.63	+1.35
AFGL 2591	-3.6(14)	4.0(10)	+24.0	+0.92	+8.67	+11.8
G327-0.6	-42.5(5)	3.3(9)	-1.96	-0.62	>15.3	-1.43
NGC 6334I	-5.7(16)	5.7(7)	-5.31	-0.79	-5.46	+15.3
G29.96	+98.2(6)	4.3(13)	-0.97	+1.10	>13.8	+19.2
G31.41	+99.1(4)	5.2(14)	>16.7	-1.68	>23.2	-1.67
G5.89	+12.0(4)	7.2(20)	-5.71	+4.85	>22.8	+70.0
G10.47	+70.0(10)	3.9(4)	-1.88	-0.68	-5.82	+4.17
<sup>a</sup>	+65.5(9)	4.3(13)	-1.63	+0.50	-6.43	+18.3
G34.26	+61.5(8)	4.9(11)	-22.1	-0.63	-8.40	-1.35
<sup>a</sup>	+57.1(4)	5.0(6)	-5.81	+2.99	-7.85	+27.4
W51e	+55.7(17)	6.3(12)	-12.5	-0.81	-10.2	+43.8
NGC 7538	-57.0(3)	3.6(10)	-4.10	+0.34	-8.02	+18.7

**Notes.** Positive fluxes denote integrated intensities ( $\int T_{\text{mb}} dV$  in K km s<sup>-1</sup>) of emission features, where upper limits denote non-detections. Negative fluxes denote integrated optical depths ( $\int \tau dV$  in km s<sup>-1</sup>) of absorption features, where lower limits denote saturated absorptions. Central velocities and FWHM line widths are in units of km s<sup>-1</sup>, and numbers in parentheses are 1 $\sigma$  errors in units of the last decimal. Line fluxes have uncertainties of  $\approx 10\%$ . <sup>(a)</sup> Source with two envelope components.

**Table A.3.** As Table A.2, for the outflows.

Source	$V_{\text{LSR}}$	$\Delta V$	1113 GHz	H <sub>2</sub> <sup>18</sup> O	1669 GHz	987 GHz
IRAS 05358	-13.13(54)	16.9(46)	+14.2	+0.42	+19.0	+23.3
IRAS 16272	-34.8(3)	6.72(3)	-1.73	n/d	-6.75	n/m
	-48.1(28)	20.4(49)	+6.16	n/d	+3.79	+11.4
NGC6334I(N)	-4.3(19)	9.2(66)	-3.58	-2.06	>14.2	n/m
	+1.0(30)	17.1(39)	+21.6	+4.33	n/m	+41.0
W43 MM1	+95.85(35)	6.65(81)	-8.25	n/d	-11.3	n/m
	+105.4(35)	10.0(48)	+5.62	n/d	+3.41	+29.9
DR21(OH)	-1.15(60)	2.17(7)	-2.61	-0.24	-2.50	n/m
	-3.9(46)	13.1(64)	+41.0	+1.21	+21.9	+80.8
W3 IRS5	-41.88(90)	3.0(14)	-0.17	+1.74	-4.57	+2.85
	-36.6(11)	21.7(73)	+26.0	n/m	+65.5	+66.3
IRAS 18089	+32.5(27)	4.6(9)	-10.8	+0.59	>17.0	n/m
	+37.3(27)	14.2(28)	+5.41	n/m	+5.67	+10.5
W33A	+29.8(15)	5.8(26)	-16.80	-0.28	-6.34	n/m
	+42.6(45)	20.6(91)	+11.1	+1.09	+10.7	+15.2
IRAS 18151	+34.3(4)	15.6(51)	+3.98	n/d	+4.21	+4.20
AFGL 2591	-10.0(17)	15.0(16)	-20.6	n/d	-38.3	+6.30
G327-0.6	-48.3(15)	6.5(26)	-6.91	-0.71	>18.44	+3.75
	-37.5(50)	15.3(47)	+2.29	+2.51	n/m	+32.6
NGC 6334I	-10.2(49)	6.3(23)	-2.21	-1.19	n/m	-1.22
	-11.0(56)	14.7(54)	-1.57	-0.29	-21.8	+55.2
G29.96	+100.7(56)	3.9(15)	-16.96	n/d	-2.98	n/m
	+99.13(39)	18.4(24)	+16.1	n/d	+24.6	+20.2
G31.41	+95.0(4)	4.0(19)	-5.36	n/d	-2.34	+6.91
	+99.65(3)	34.0(3)	n/m	n/d	n/m	+12.8
G5.89	+44.1(30)	14.8(4)	+7.55	n/m	+18.46	n/m
	+27.7(32)	19.3(65)	+20.5	+3.86	+35.71	+119.4
	-9.4(28)	30(13)	-3.78	-1.64	+34.14	+65.7
	-26.9(5)	13.9(45)	-0.70	-0.62	n/m	n/m
	-49.5(85)	16.2(72)	-1.14	-0.30	-1.53	+4.83
G10.47	+62.0(5)	4.1(5)	+2.70	n/m	n/m	+4.45
	+80.2(51)	8.3(6)	+4.71	n/m	-1.36	+5.47
G34.26	+53.7(77)	4.2(13)	-9.60	+1.20	-9.14	n/m
	+57.0(85)	18.1(81)	-1.66	-0.96	-10.7	+25.8
W51e	+67.7(13)	6.9(28)	-14.3	-0.25	-28.2	-1.44
	+61.7(13)	14(11)	-17.6	-0.91	-4.70	+62.5
NGC 7538 IRS1	-57.7(51)	2.25(43)	-2.36	-0.13	-0.80	n/m
	-56.8(58)	13.4(24)	+10.3	+1.67	+2.66	+17.4

**Notes.** The symbol n/d means that the outflow is not at all detected in this line; the symbol n/m means that only this component has no match in this line.

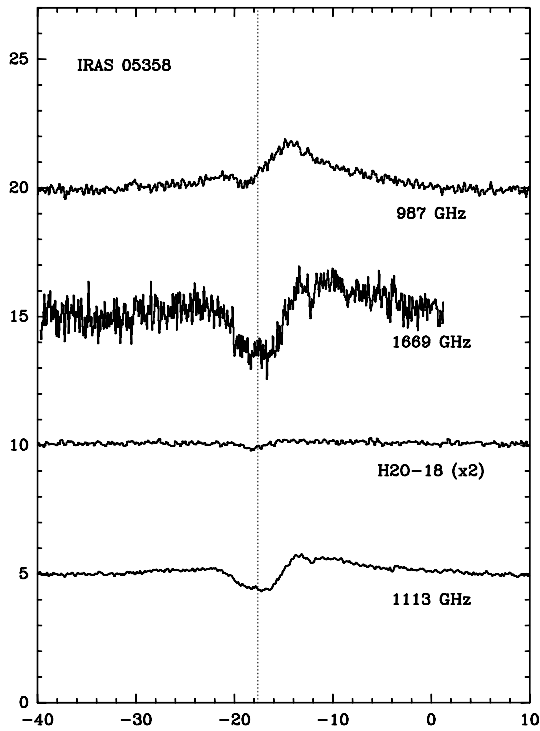


**Table A.4.** As Table A.2, for the foreground clouds.

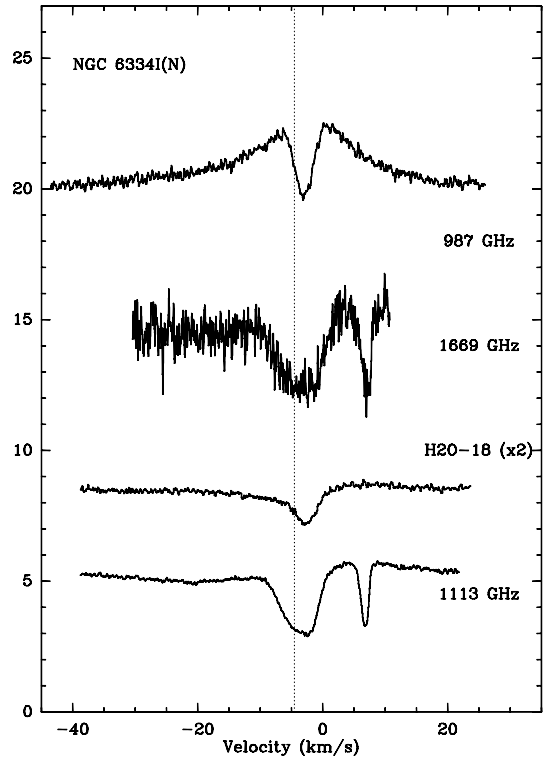
Source	$V_{\text{LSR}}$	$\Delta V$	1113 GHz	1669 GHz	$N(\text{H}_2\text{O})$	o/p
IRAS 16272	-41.38(8)	1.67(3)	-2.16	>6.33	>35	>5.9
	-38.80(23)	1.94(5)	-1.36	-2.50	15	3.7
NGC6334I(N)	+6.86(13)	1.49(12)	-1.52	>5.12	>27	>6.6
W43 MM1	+62.1(6)	1.04(8)	-0.094	-0.11	0.73	2.4
	+67.21(4)	1.81(8)	-1.13	-1.80	11	3.2
	+70.96(8)	1.6(2)	-0.73	-1.15	7.1	3.2
	+77.(7)	4.1(6)	-1.63	-1.11	9.0	1.4
	+79.7(1)	1.93(5)	-1.48	-3.67	21	5.0
	+82.4(2)	2.49(5)	-1.63	-2.58	16	3.2
	+87.83(12)	2.1(4)	-0.20	-0.49	2.75	4.9
DR21(OH)	+14.9(1)	1.5(2)	-0.35	-0.90	5.0	5.1
	+12.3(1)	1.8(2)	-0.47	-2.03	11	8.6
	+10.0(1)	1.7(7)	-0.51	-2.53	13	9.9
	+8.4(4)	1.9(2)	-3.78	-0.80	13	0.4
	+6.9(1)	1.22(1)	-0.51	-0.38	3.0	1.5
W3 IRS5	-20.8(1)	1.04(8)	-0.61	>3.56	>18	>12
IRAS 18089	+17.6(1)	3.4(6)	-2.81	-7.24	40.3	5.2
	+19.7(3)	2.8(11)	-0.91	-2.34	13.0	5.1
	+21.1(1)	2.2(4)	-0.95	-1.42	8.8	3.0
	+23.9(2)	1.9(1)	-0.82	-2.68	14.4	6.5
	+25.7(2)	2.67(7)	-1.46	-4.00	22.0	5.5
W33A	+24.01(7)	1.5(1)	-0.66	-1.26	7.4	3.8
IRAS 18151	+27.7(2)	0.7(2)	-0.36	-1.06	5.8	5.9
AFGL 2591	-0.67(13)	3.0(6)	>9.54	>7.63	>58	~1.6
	+13 <sup>a</sup>					
G327-0.6 <sup>d</sup>	-38.02(4)	1.18(33)	-0.06	-0.19	1.1	6.0
NGC 6334I	+0.48(11)	2.58(52)	-1.45	-3.42	13.8	2.7
	+6.63(17)	1.61(16)	-2.11	>5.63	>31.2	>5.3
	+8.72(39)	1.27(52)	-0.20	-0.17	1.3	1.7
G29.96	+103.5(1)	1.7(1)	-1.09	-2.91	16	5.3
	+91.6(2)	3.4(5)	-0.67	-0.59	4.3	1.8
G31.41	+83.0(3)	3.4(6)	-0.48	-1.27	7.0	5.3
G5.89	+19.85(1)	4.4(6)	-1.90	-25.17 <sup>b</sup>	121.7 <sup>b</sup>	26.5 <sup>b</sup>
G10.47	+42.3(1)	1.1(3)	-0.16	-0.24	1.5	3.0
	+47.06(4)	1.35(7)	-0.09	-0.19	1.1	4.5
	+84.9(1)	2.6(10)	-0.19	-1.67 <sup>b</sup>	8.2 <sup>b</sup>	17.7 <sup>b</sup>
	+91.21(6)	1.4(4)	-0.08	-0.32	1.7	8.3
G34.26	+11.4(1)	2.6(1)	-6.43	-14.0	80	4.4
	+14.26(3)	1.18(1)	-0.35	-0.55	3.4	3.1
	+27.09(1)	1.6(3)	-1.52	-5.48	29	7.2
W51e	+48.66(1)	2.27(1)	-0.42	-1.14	6.3	5.4
	+45.57(9)	1.2(1)	-2.51	>4.42	>26	>3.5
	+12.7 <sup>c</sup>					
	+6.48(5)	2.4(2)	-2.63	-2.99	20	2.3
	+4.87(1)	1.1(2)	-0.27	-0.84	4.5	6.2

**Notes.** The last two columns list the derived  $\text{H}_2\text{O}$  column density (in  $10^{12} \text{ cm}^{-2}$ ) and ortho/para ratio. <sup>(a)</sup> Blended with  $\text{H}_2\text{O}$  1661 GHz line emission from other sideband; see Choi et al. (in prep.). <sup>(b)</sup> Highly uncertain because 1669 GHz line is saturated. <sup>(c)</sup> Only detected at 1113 GHz. <sup>(d)</sup> WBS spectra of this source show spatially extended  $\text{H}_2\text{O}$  foreground absorbers between  $-20$  and  $-3 \text{ km s}^{-1}$  (Leurini, priv. comm.).

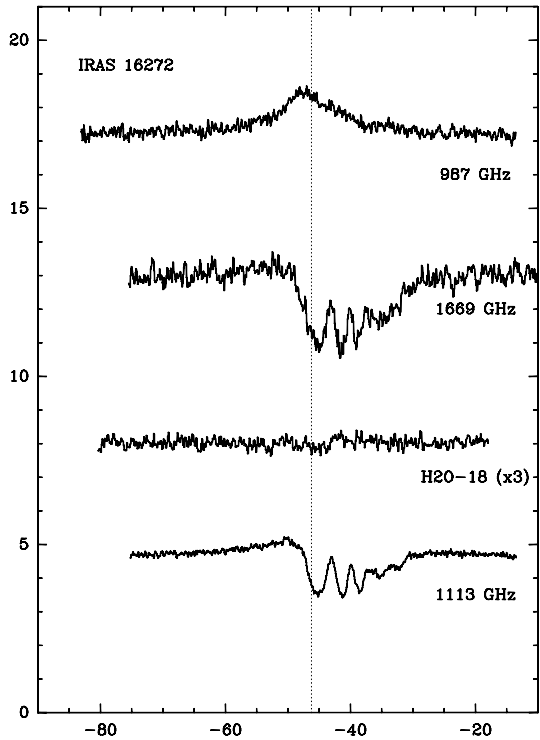
Appendix B: Observed line profiles



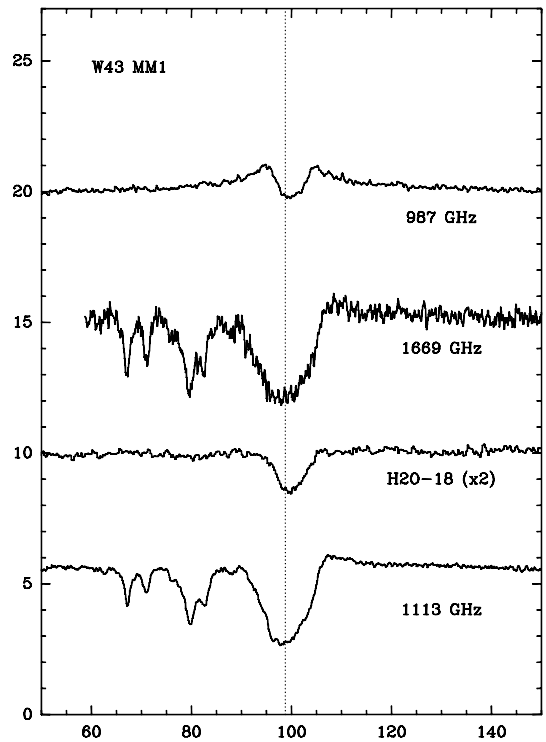
**Fig. B.1.** Spectra of the H<sub>2</sub>O  $1_{11}-0_{00}$  (bottom), H<sub>2</sub><sup>18</sup>O  $1_{11}-0_{00}$  (second from bottom), H<sub>2</sub>O  $2_{12}-1_{01}$  (third from bottom) and H<sub>2</sub>O  $2_{02}-1_{11}$  (top) lines toward IRAS 05358. The vertical scale is  $T_{\text{mb}}$  in K, and the dashed line denotes the systemic velocity determined from ground-based observations. The spectra have been shifted vertically for clarity.



**Fig. B.3.** As Fig. B.1, toward NGC 6334I(N).



**Fig. B.2.** As Fig. B.1, toward IRAS 16272.



**Fig. B.4.** As Fig. B.1, toward W43 MM1.

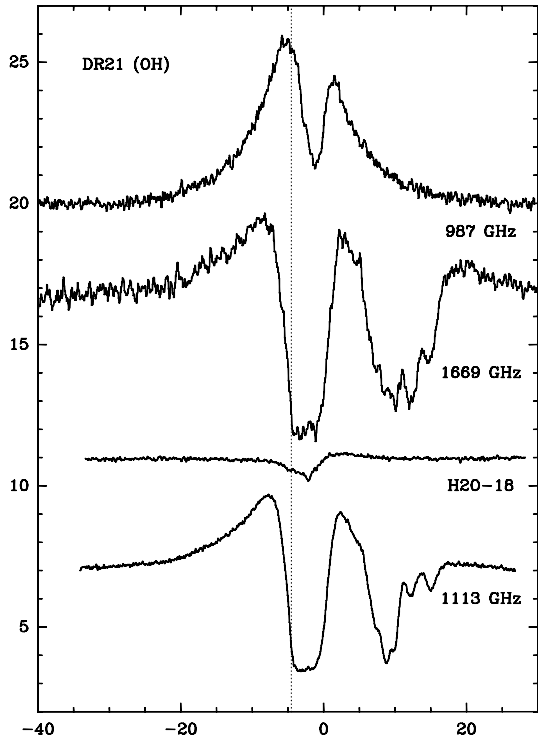


Fig. B.5. As Fig. B.1, toward DR21(OH).

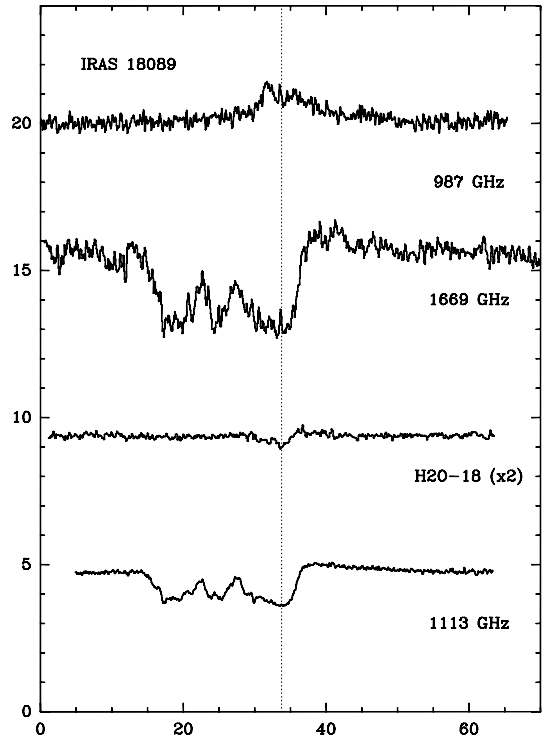


Fig. B.7. As Fig. B.1, toward IRAS 18089.

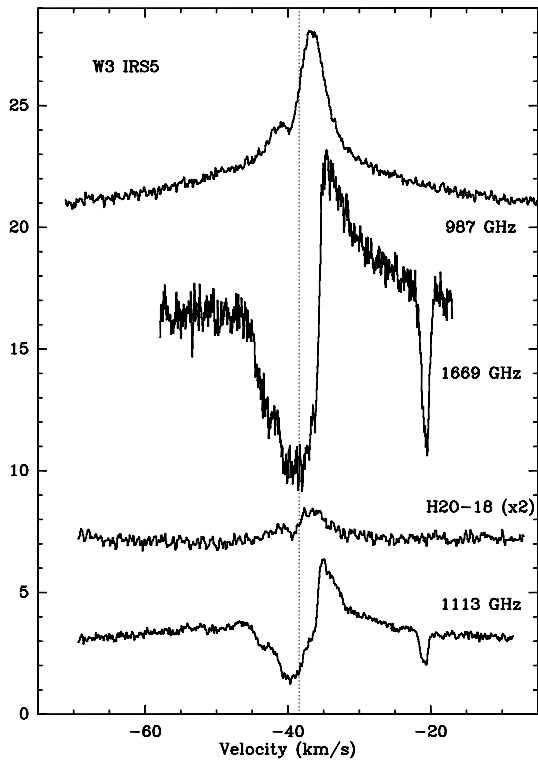


Fig. B.6. As Fig. B.1, toward W3 IRS5.

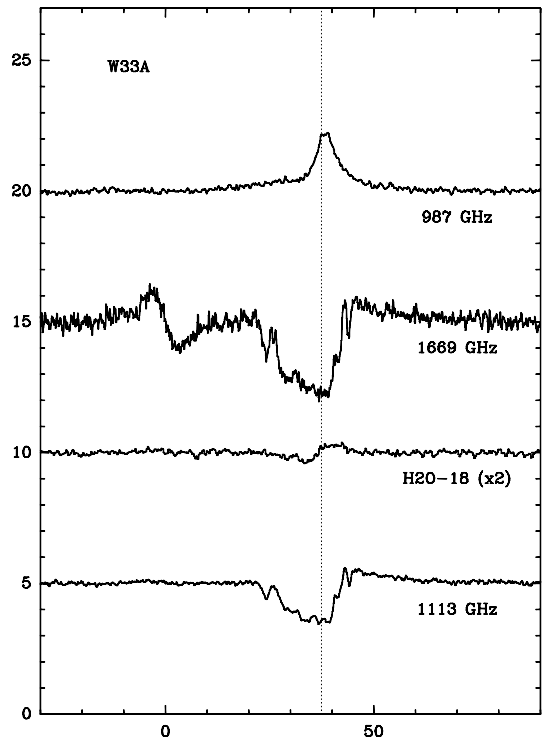
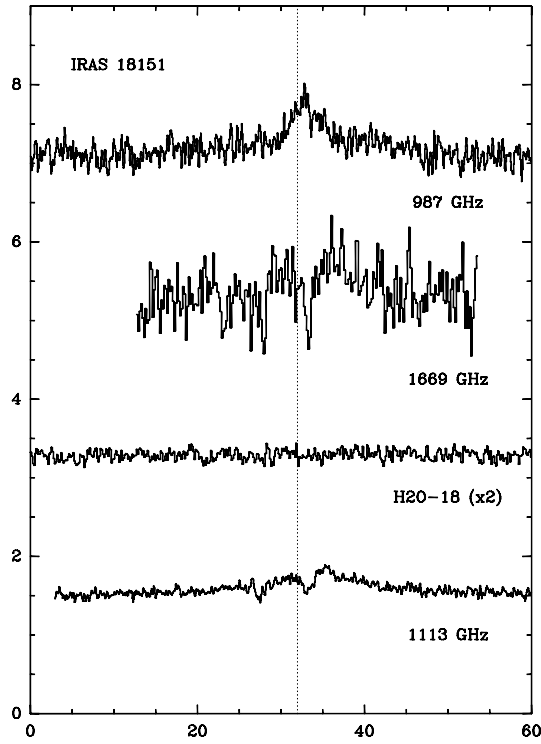
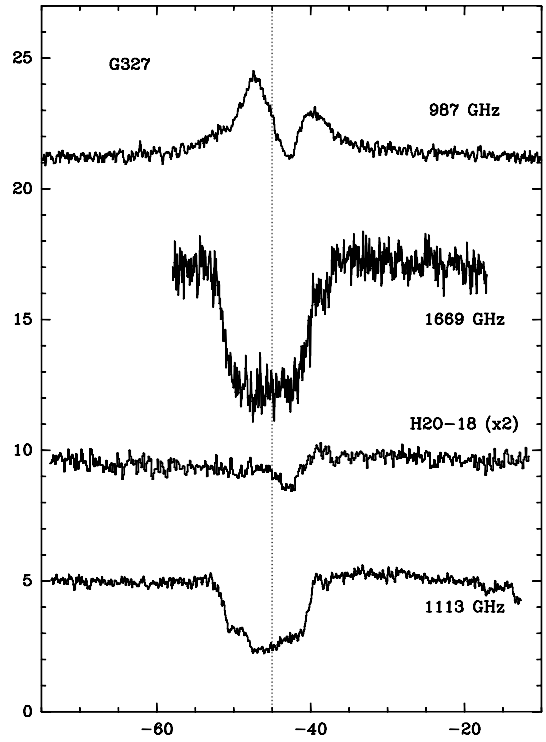


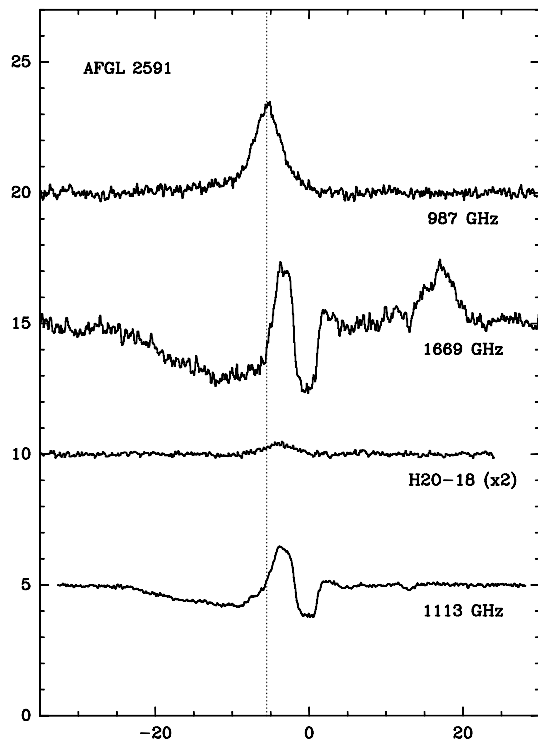
Fig. B.8. As Fig. B.1, toward W33A. The emission/absorption feature near  $0 \text{ km s}^{-1}$  in the 1669 GHz spectrum is the H<sub>2</sub>O 1661 GHz line from the other sideband.



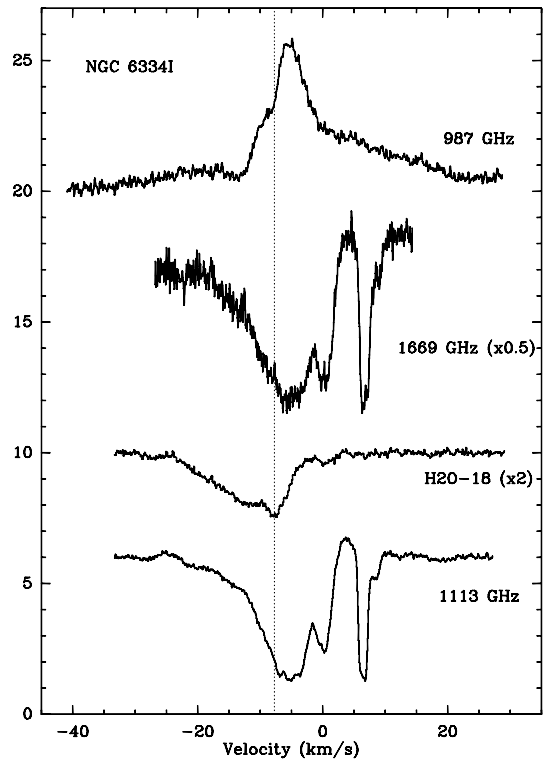
**Fig. B.9.** As Fig. B.1, toward IRAS 18151.



**Fig. B.11.** As Fig. B.1, toward G 327.



**Fig. B.10.** As Fig. B.1, toward AFGL 2591. The emission feature near  $+18 \text{ km s}^{-1}$  in the 1669 GHz spectrum is the H<sub>2</sub>O 1661 GHz line from the other sideband.



**Fig. B.12.** As Fig. B.1, toward NGC 6334I.

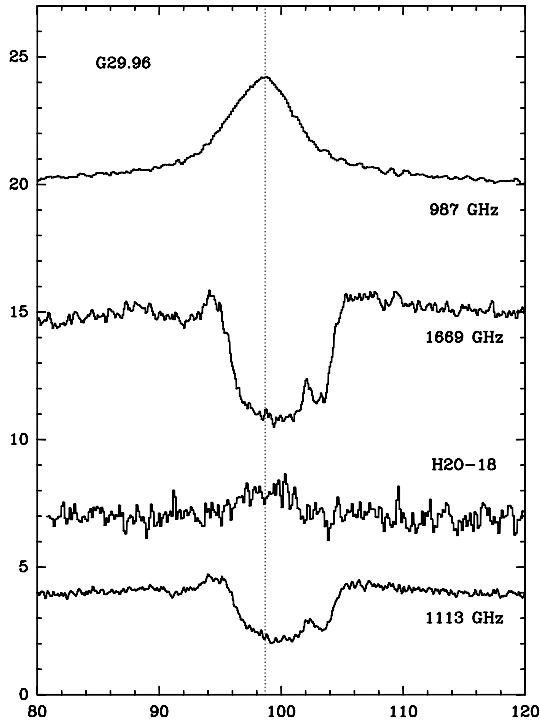


Fig. B.13. As Fig. B.1, toward G29.96.

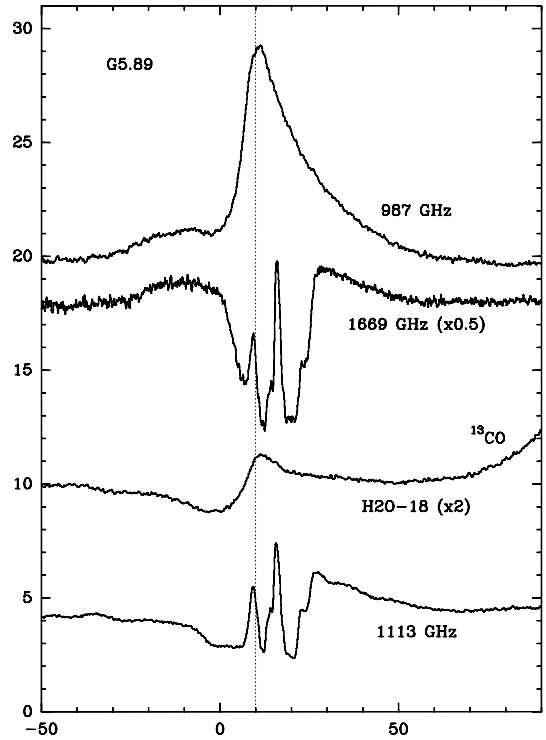


Fig. B.15. As Fig. B.1, toward G5.89. The emission feature near +90 km s<sup>-1</sup> in the H<sub>2</sub><sup>18</sup>O spectrum is the <sup>13</sup>CO 10–9 line.

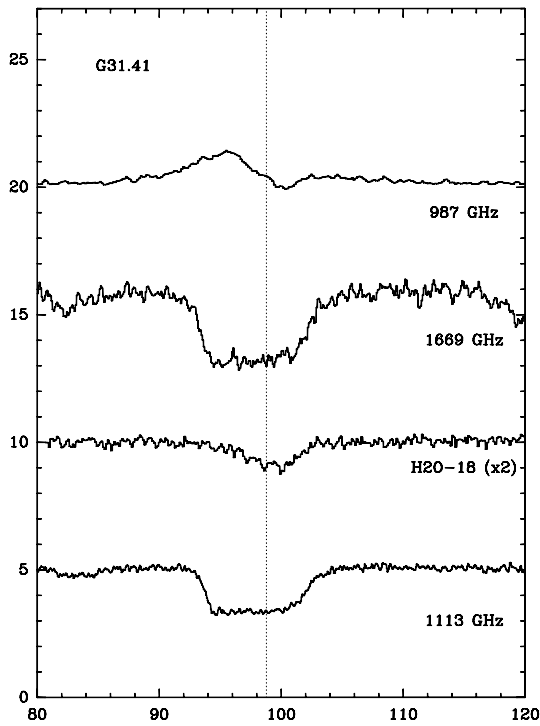


Fig. B.14. As Fig. B.1, toward G31.41.

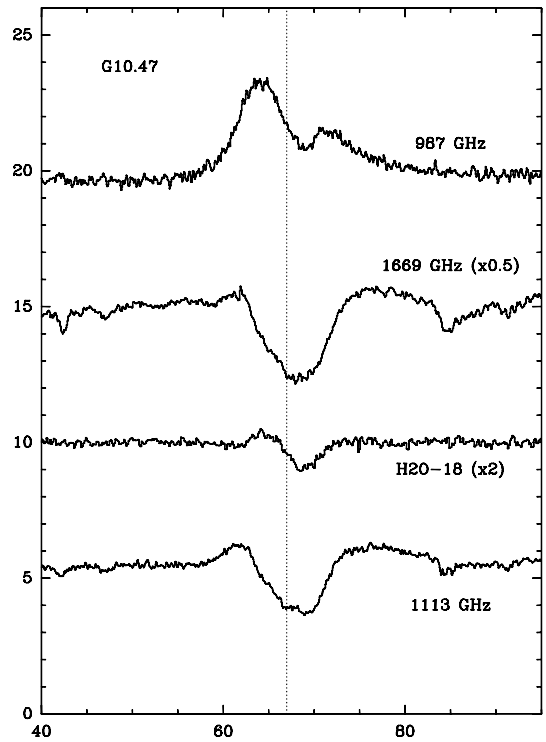


Fig. B.16. As Fig. B.1, toward G10.47.

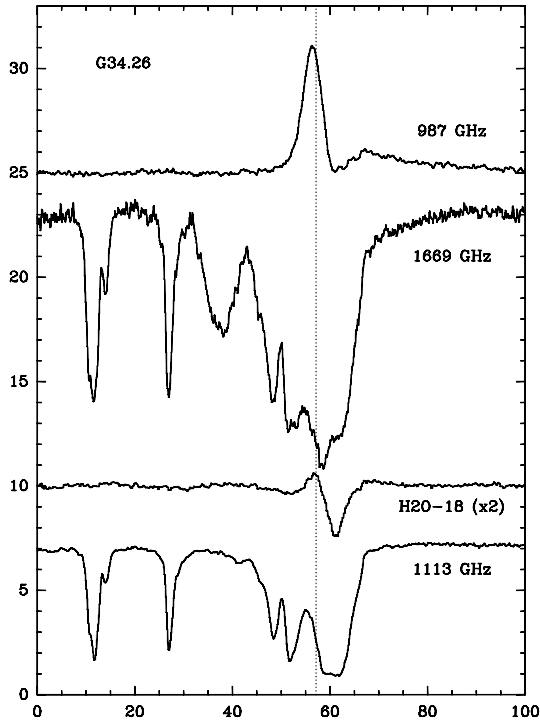


Fig. B.17. As Fig. B.1, toward G34.26.

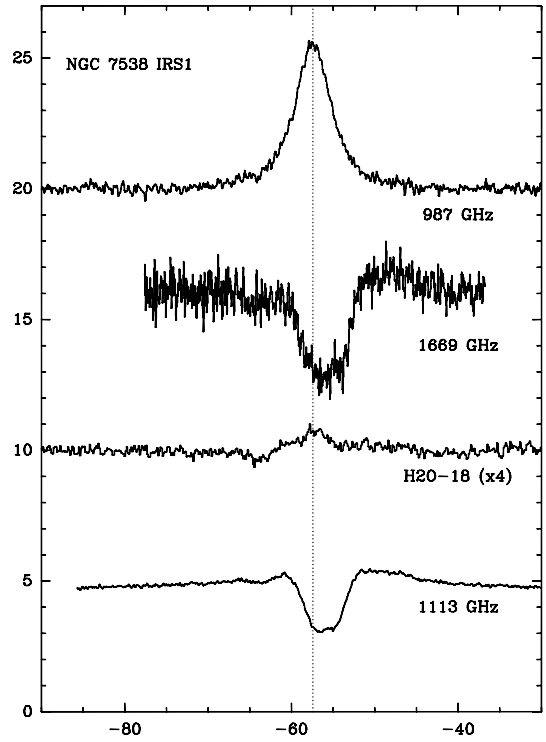


Fig. B.19. As Fig. B.1, toward NGC 7538 IRS1.

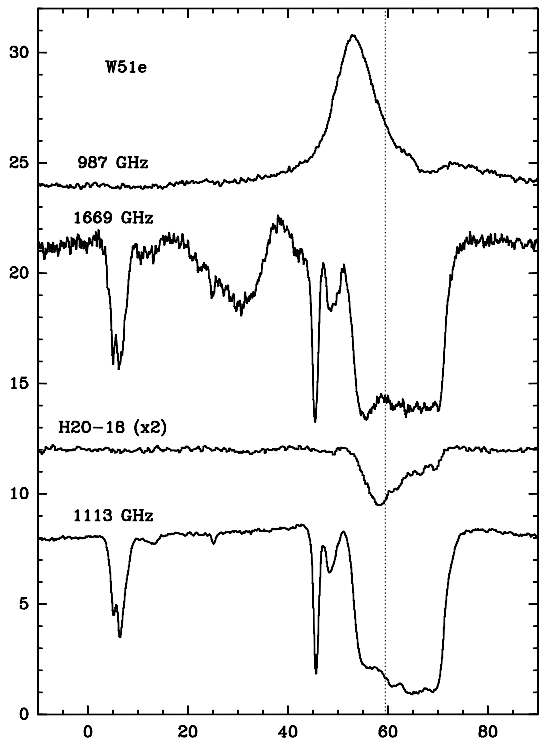
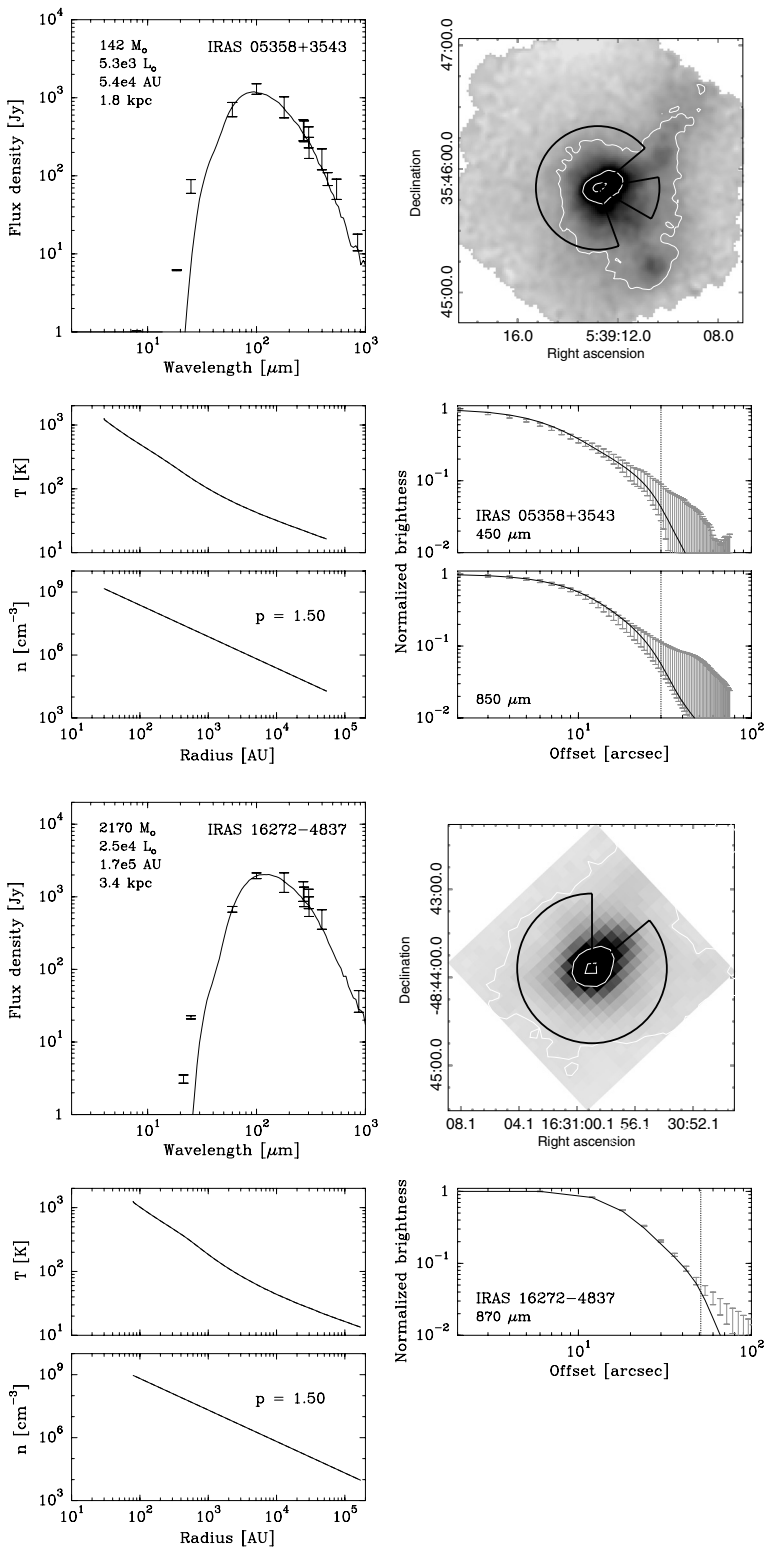


Fig. B.18. As Fig. B.1, toward W51e.

Appendix C: Continuum models



**Fig. C.1.** Continuum model for IRAS 05358. *Top left:* spectral energy distribution. *Top right:* submillimeter image with  $3\sigma$  contour marked in white and useable map sectors marked in black. *Bottom left:* temperature and density structure as a function of radius. *Bottom right:* submillimeter emission profiles. The numbers in the top left panel are the modeled envelope mass, the modeled luminosity, the adopted envelope size and the adopted distance.

**Fig. C.2.** As Fig. C.1, for IRAS 16272.

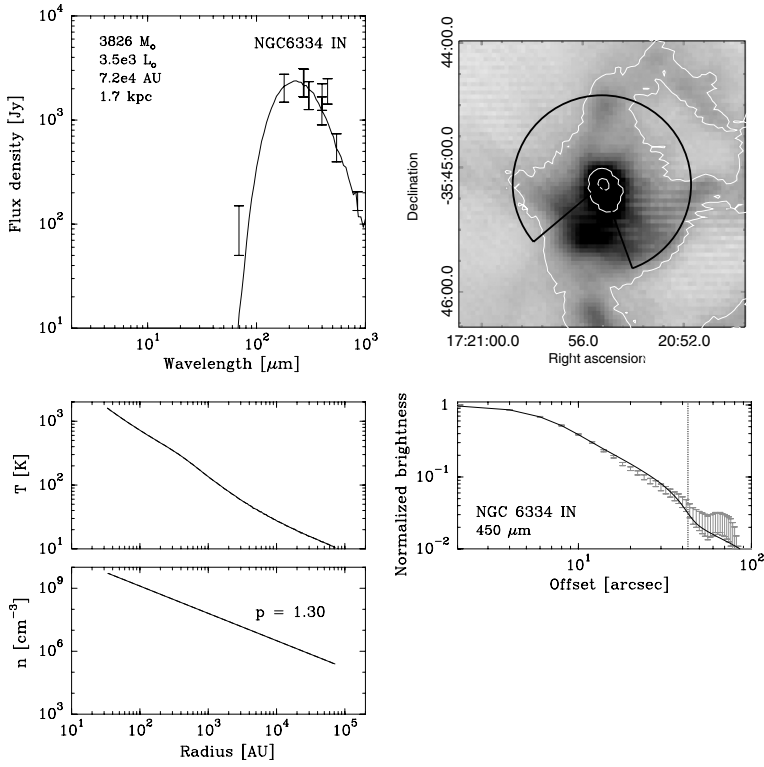


Fig. C.3. As Fig. C.1, for NGC 6334I(N).

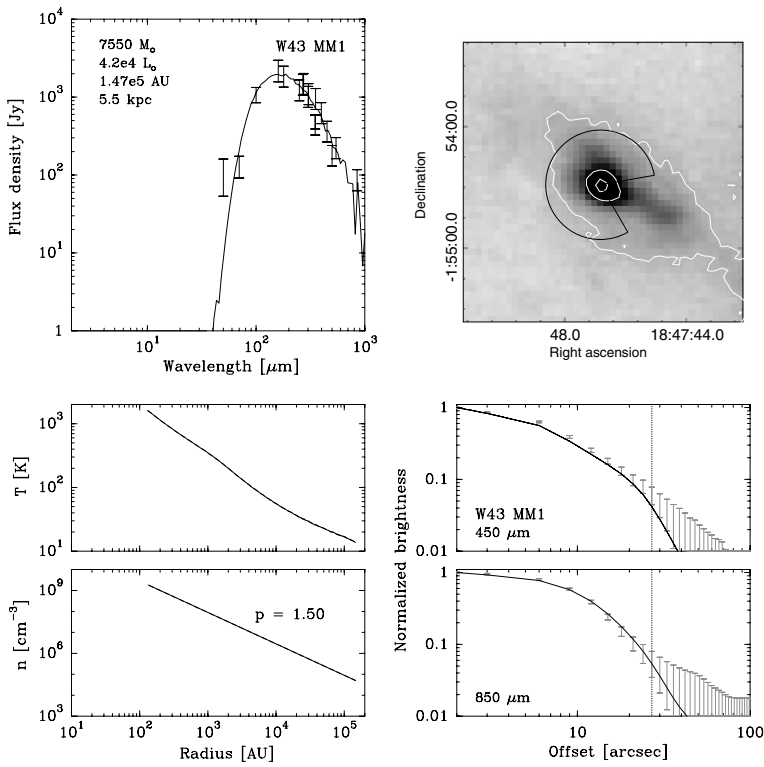


Fig. C.4. As Fig. C.1, for W43-MM1.



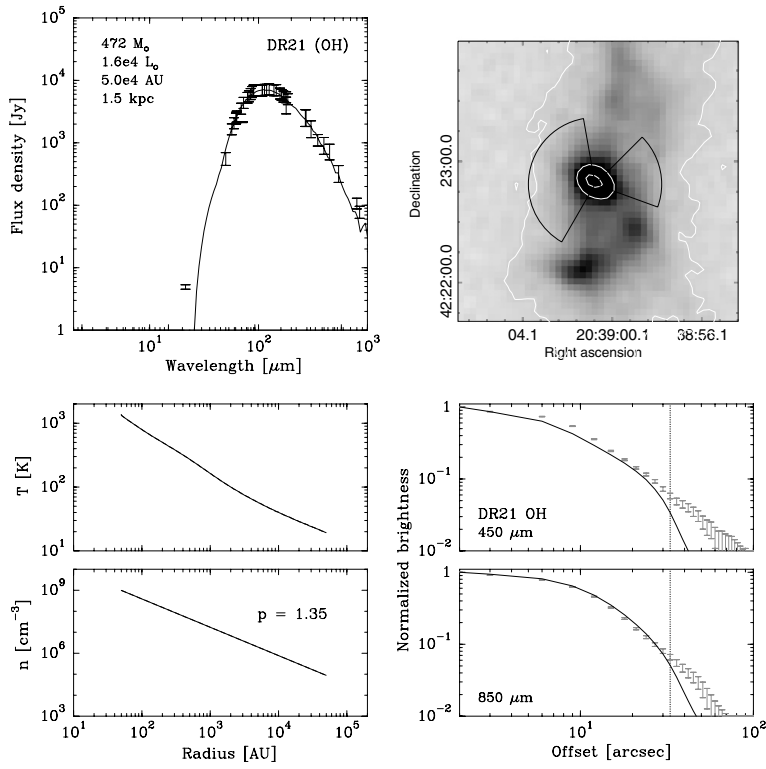


Fig. C.5. As Fig. C.1, for DR21(OH).

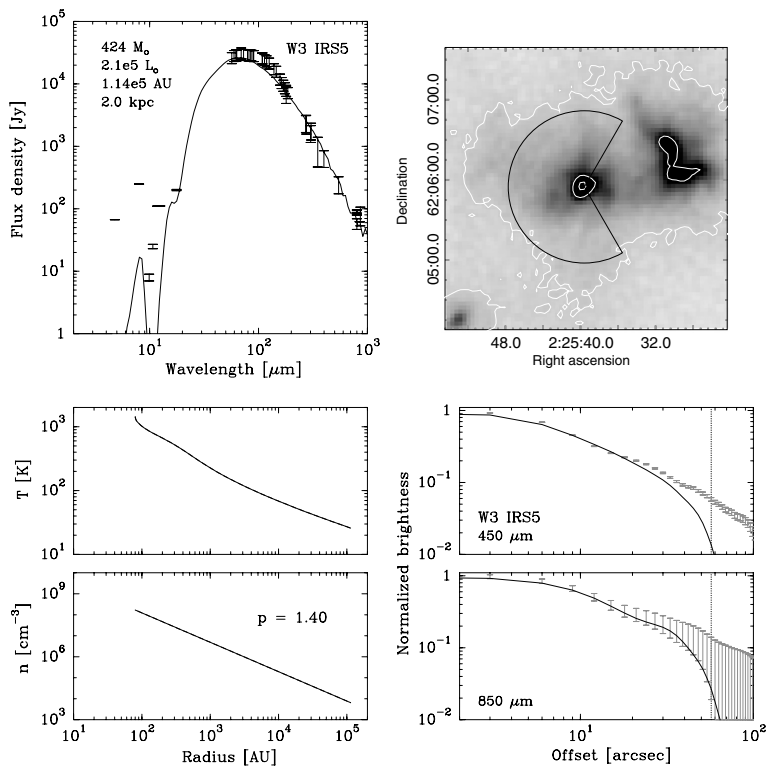


Fig. C.6. As Fig. C.1, for W3 IRS5.

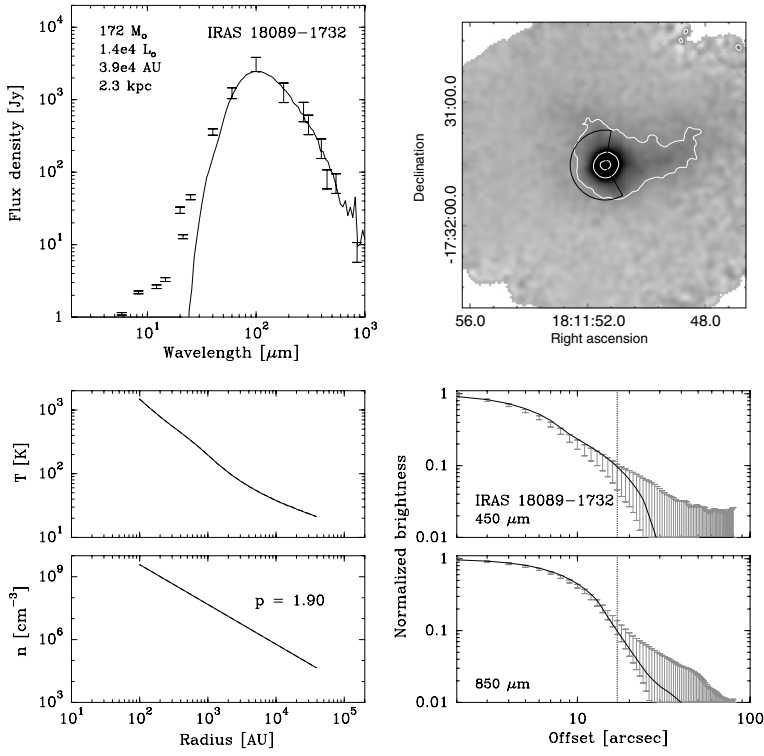


Fig. C.7. As Fig. C.1, for IRAS 18089.

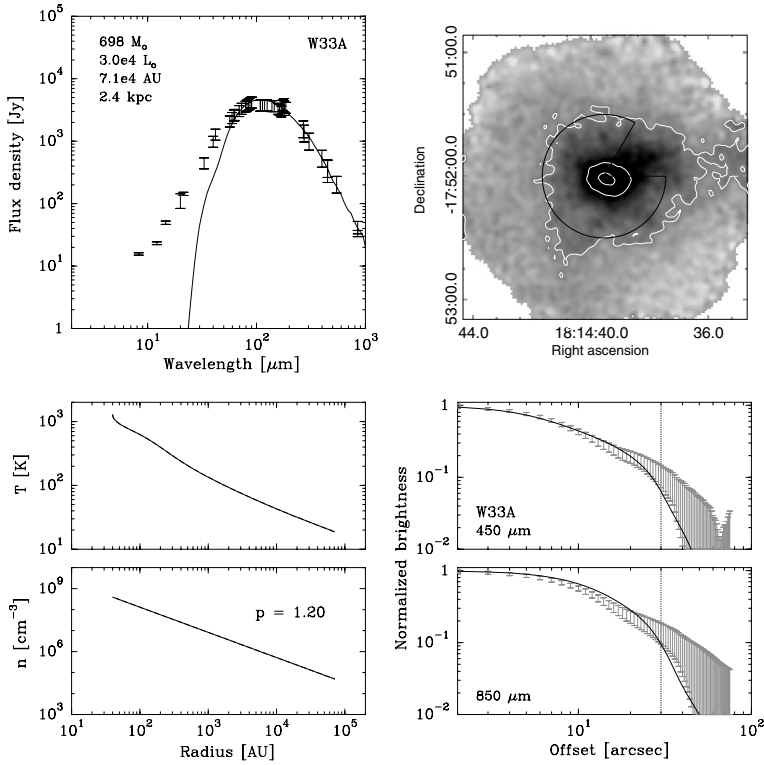


Fig. C.8. As Fig. C.1, for W33A.

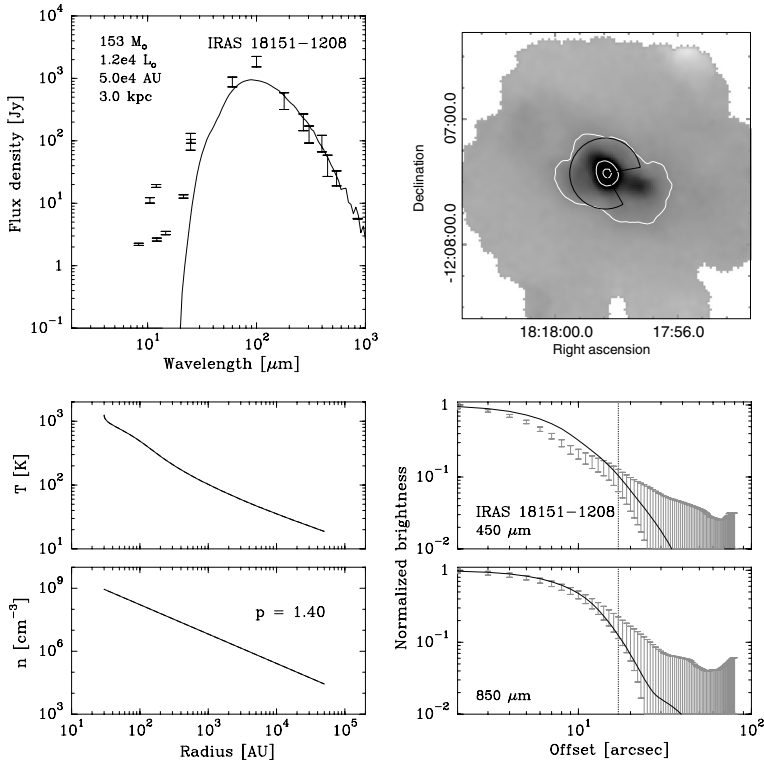


Fig. C.9. As Fig. C.1, for IRAS 18151.

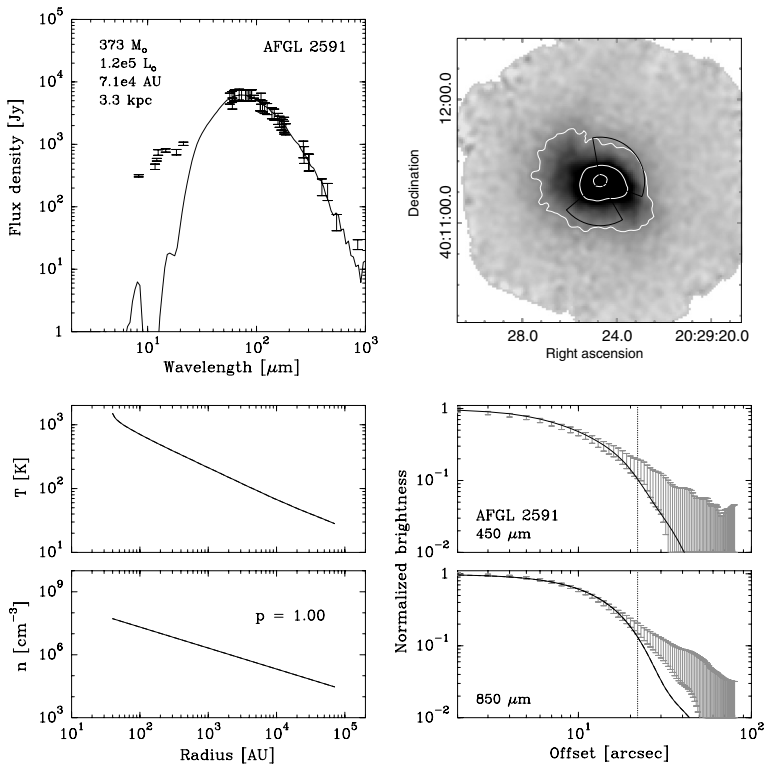


Fig. C.10. As Fig. C.1, for AFGL 2591.

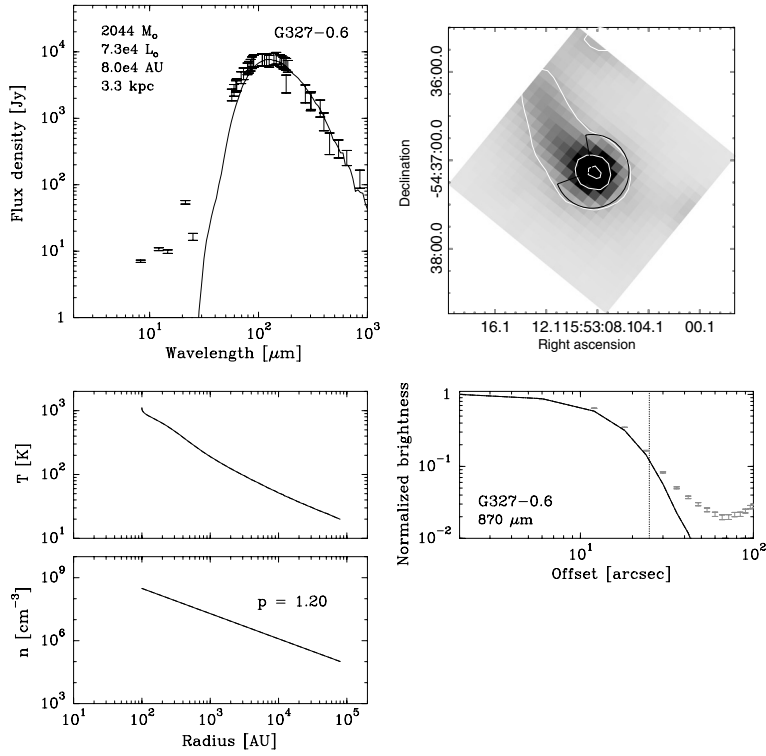


Fig. C.11. As Fig. C.1, for G327-0.6.

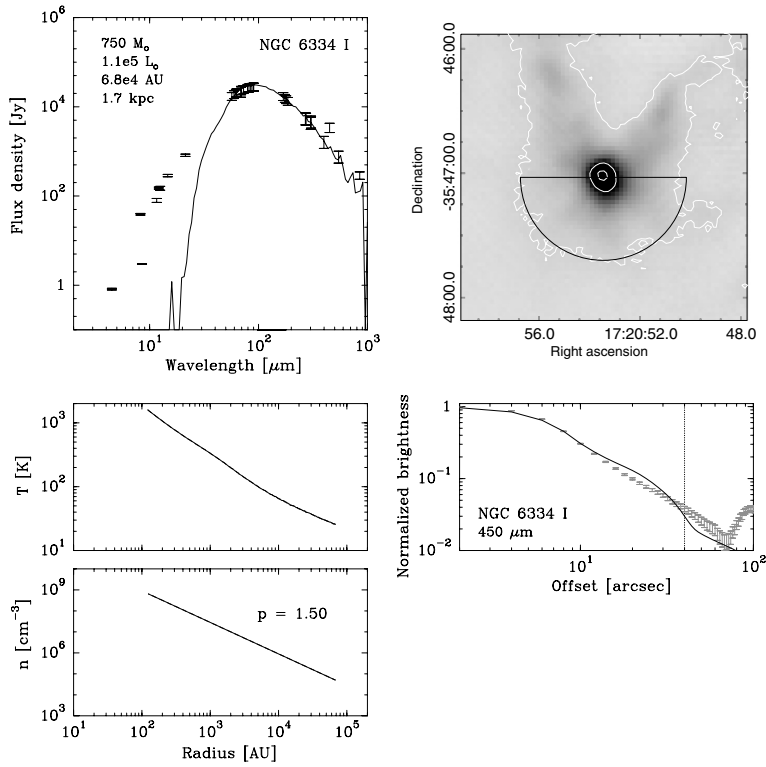


Fig. C.12. As Fig. C.1, for NGC 6334I.

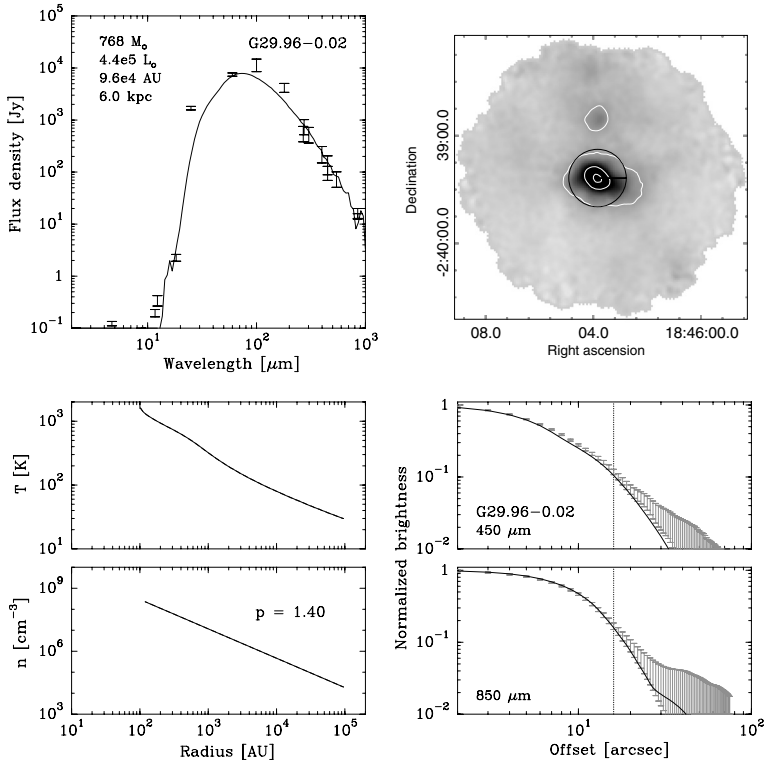


Fig. C.13. As Fig. C.1, for G29.96.

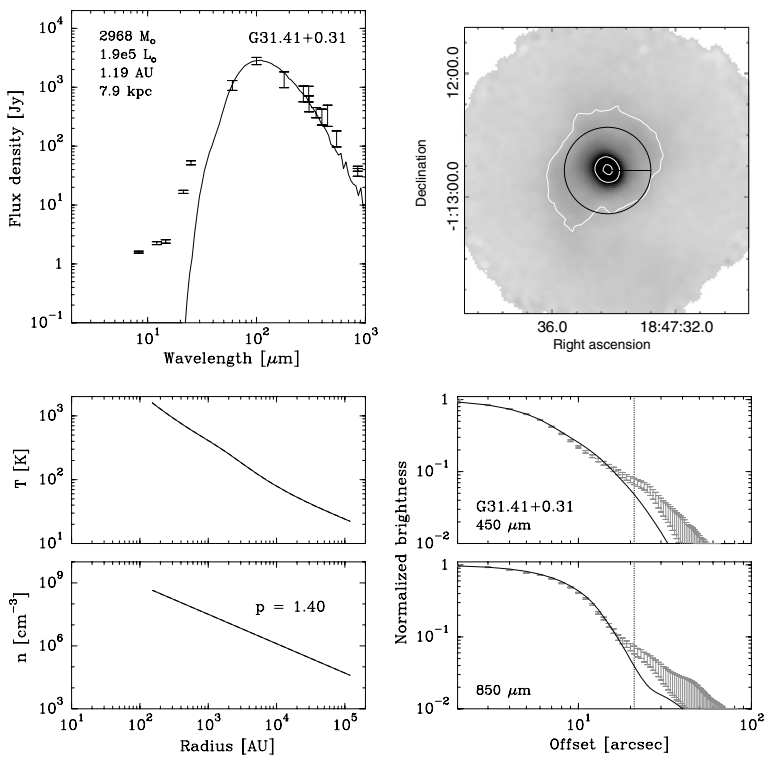


Fig. C.14. As Fig. C.1, for G31.41.

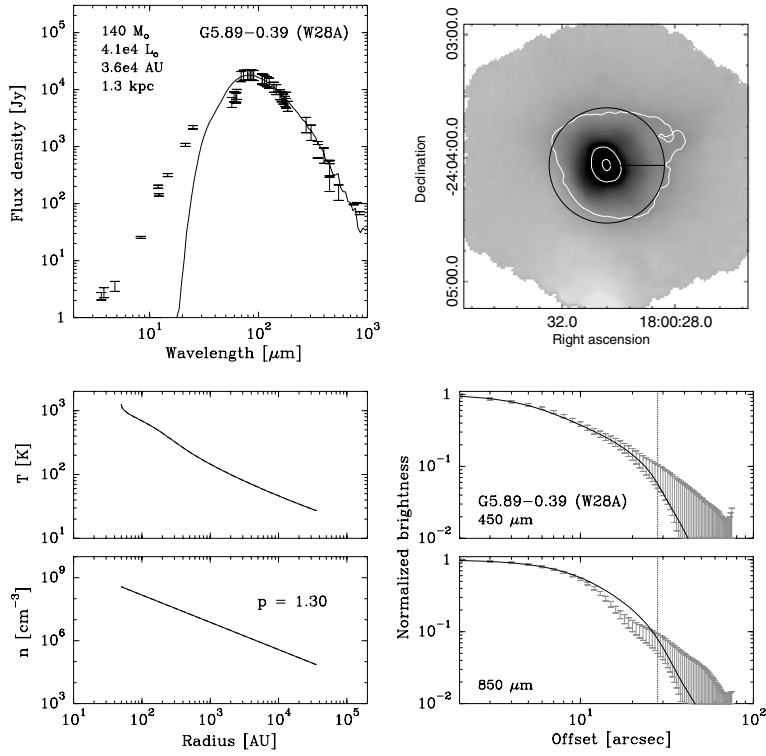


Fig. C.15. As Fig. C.1, for G5.89.

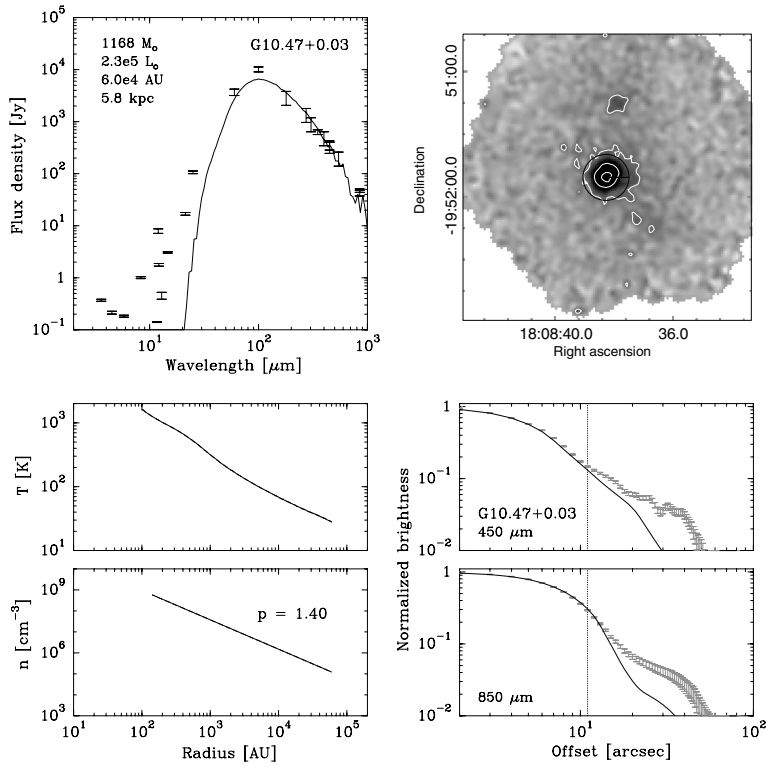


Fig. C.16. As Fig. C.1, for G10.47.

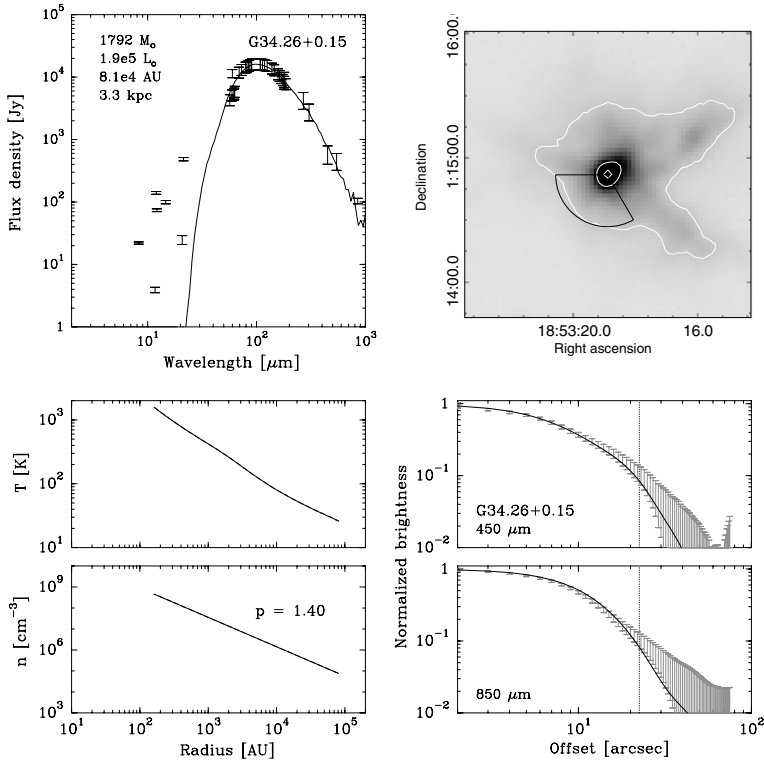


Fig. C.17. As Fig. C.1, for G34.26.

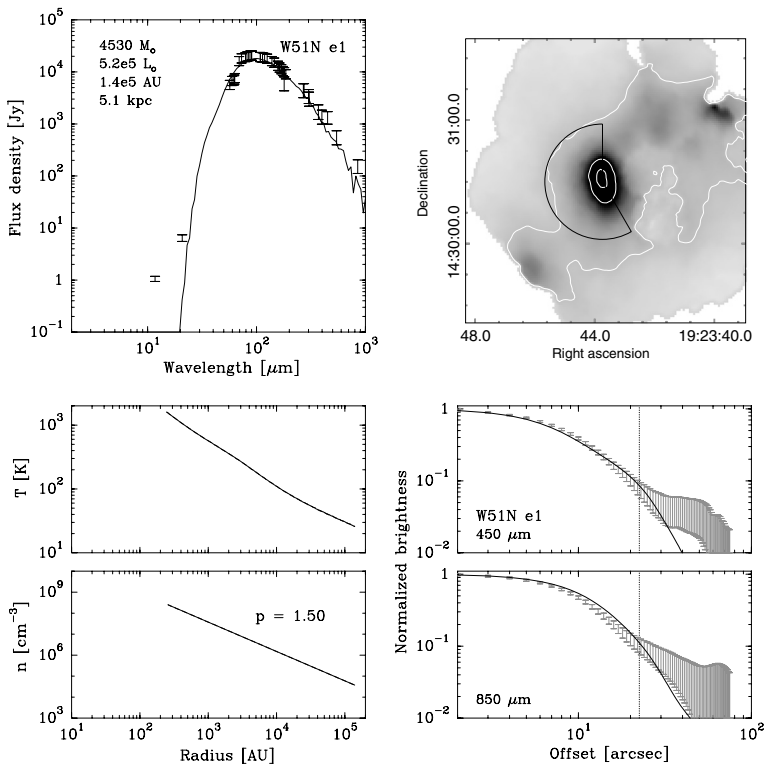
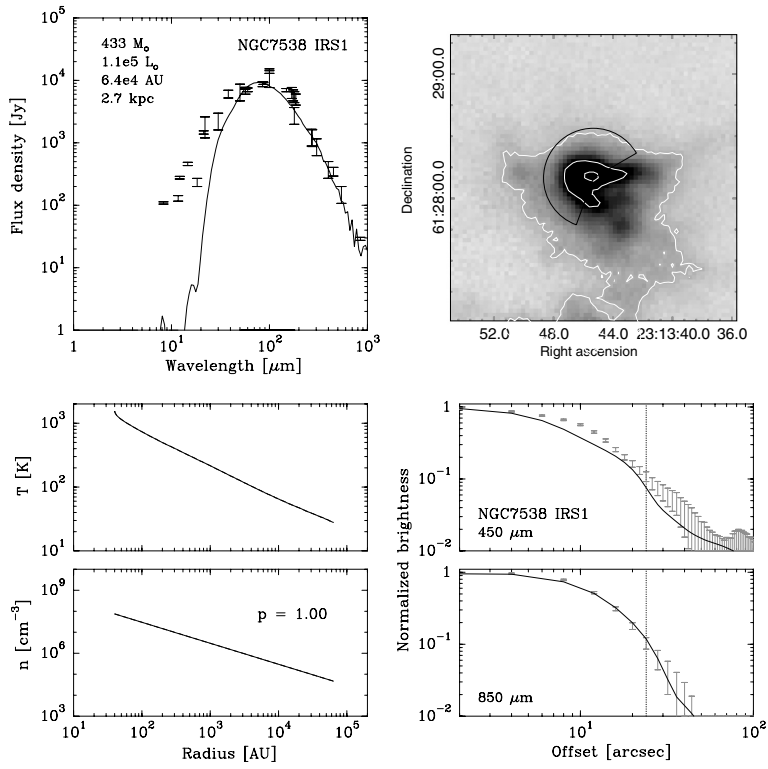


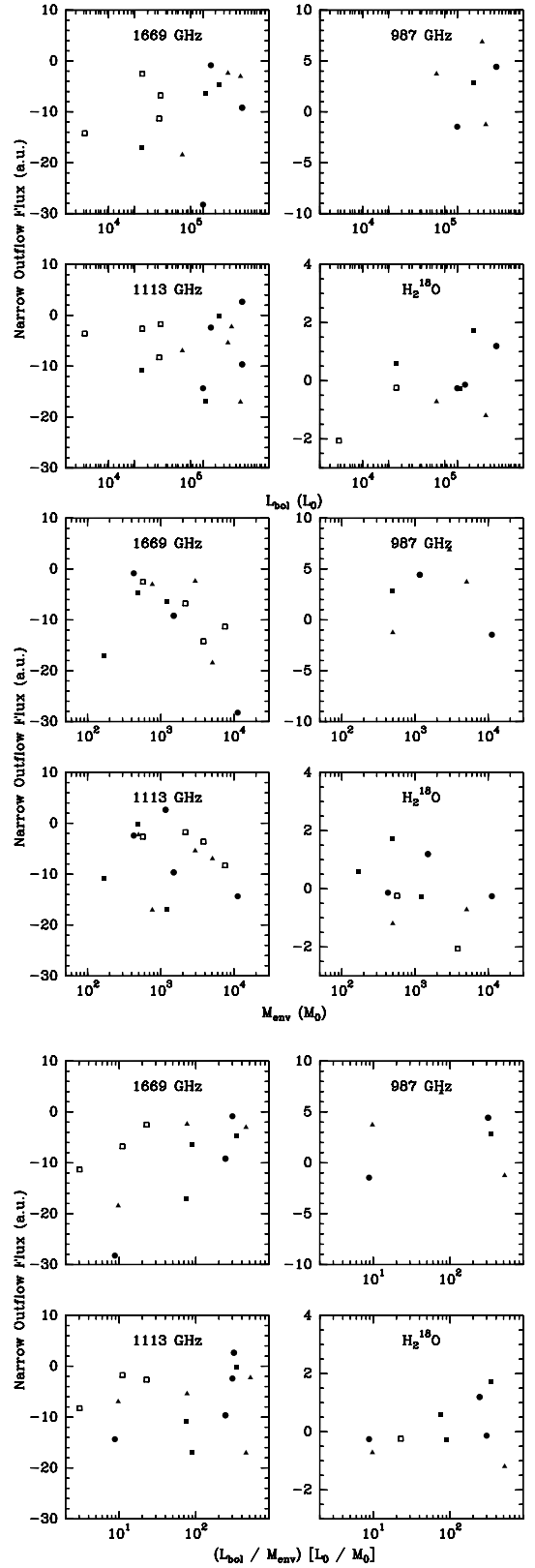
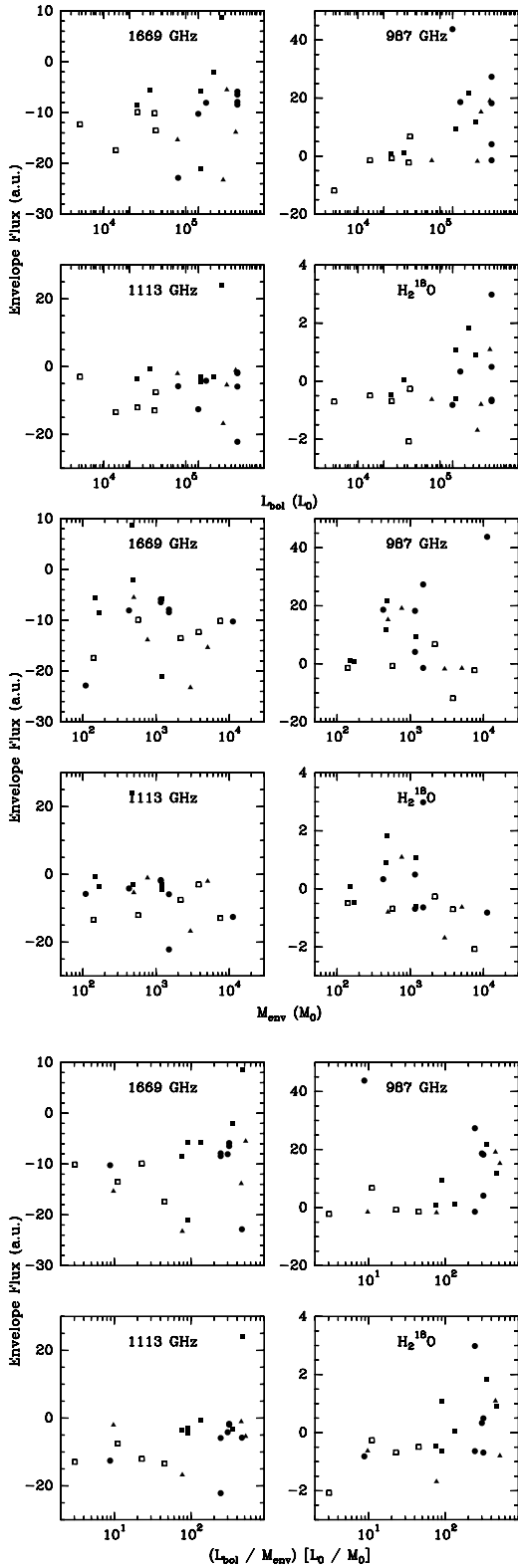
Fig. C.18. As Fig. C.1, for W51e.



**Fig. C.19.** As Fig. C.1, for NGC 7538 IRS1.

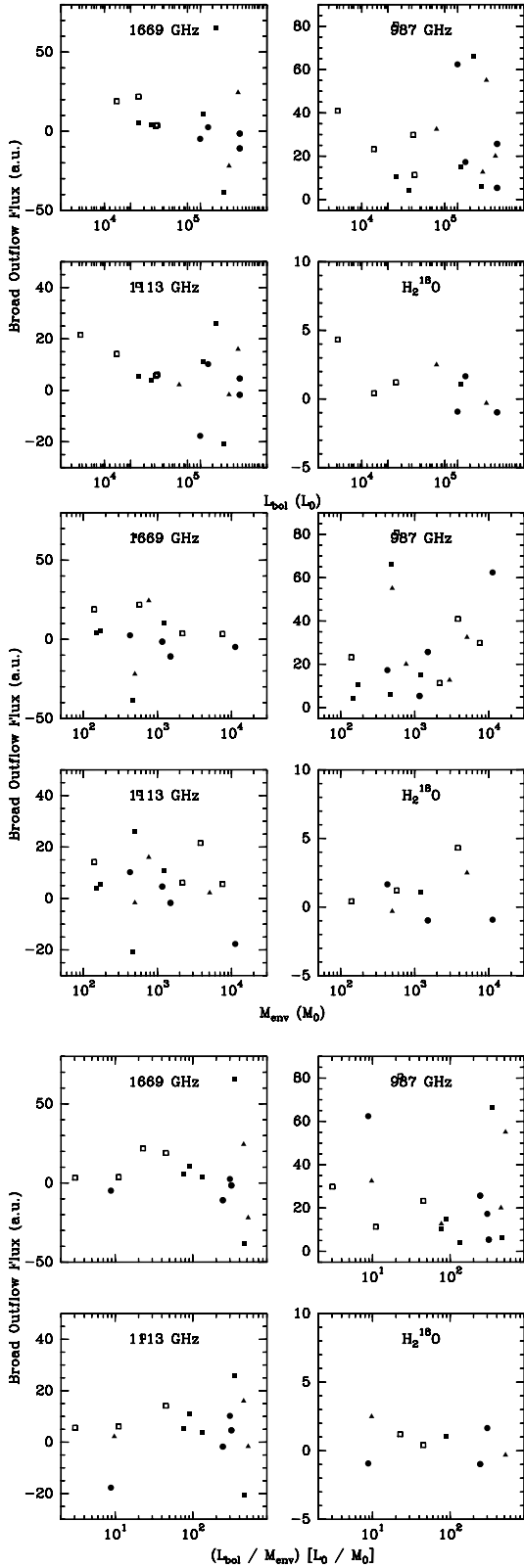


## Appendix D: Correlation analysis

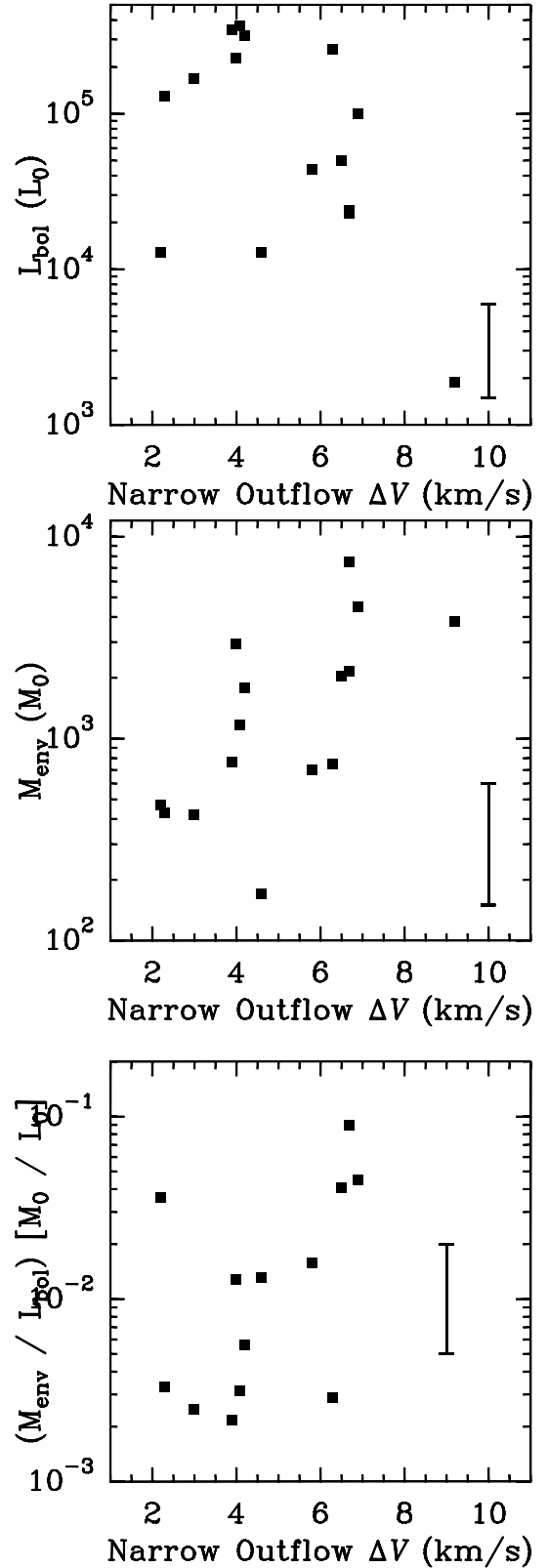


**Fig. D.1.** Observed line fluxes (positive: emission) or absorptions (negative: absorption) versus bolometric luminosity  $L$  of the source (*top*), versus envelope mass  $M$  (*middle*), and versus the ratio  $L/M$  (*bottom*). Open squares: mid-IR-quiet HMPOs; filled squares: mid-IR-bright HMPOs; triangles: hot molecular cores; circles: ultracompact HII regions. Uncertainties in line fluxes are  $\approx 10\%$ , while the typical uncertainty of masses, luminosities and  $L/M$  ratios is a factor of 2.

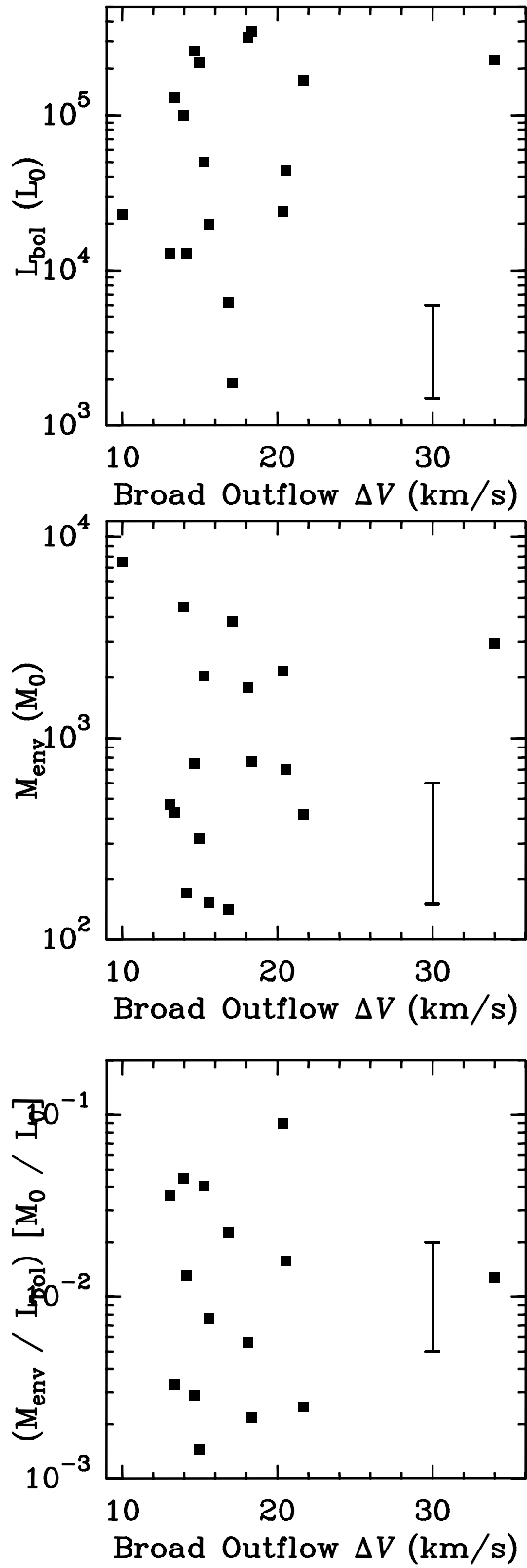
**Fig. D.2.** As Fig. D.1, for the narrow outflow component.



**Fig. D.3.** As Fig. D.1, for the broad outflow component.



**Fig. D.4.** Line width of the narrow outflow component versus bolometric luminosity  $L$  of the source (*top*), versus envelope mass  $M$  (*middle*), and versus the ratio  $L/M$  (*bottom*). The error bar denotes the typical uncertainty in mass, luminosity, and  $L/M$  ratio; the uncertainty in line width is smaller than the symbol size.



**Fig. D.5.** As Fig. D.4, for the broad outflow component.