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Invariant Relative Orbits Taking into Account Third-Body Perturbation

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ABSTRACT

For a satellite in an orbit of more than 1600 km in altitude, the effects of Sun and Moon on the orbit can't be negligible. Working with mean orbital elements, the secular drift of the longitude of the ascending node and the sum of the argument of perigee and mean anomaly are set equal between two neighboring orbits to negate the separation over time due to the potential of the Earth and the third body effect. The expressions for the second order conditions that guarantee that the drift rates of two neighboring orbits are equal on the average are derived. To this end, the Hamiltonian was developed. The expressions for the non-vanishing time rate of change of canonical elements are obtained.

Keywords: Invariant Relative Orbits; Third-Body Perturbation; Hamiltonian

1. Introduction

Formation flying is a key technology enabling a number of missions which a single satellite cannot accomplish: from remote sensing to astronomy. The relative motion, which shows no drift even in presence of a large disturbance, could be a very attractive solution. To maintain the formation and constellation, the relative drifts due to the perturbation between the spacecraft should be carefully considered. Invariant Relative Orbits shows no drift between the spacecraft due to the perturbation even if in presence of a large disturbance.

The literature is wealth with works dealing with designing certain invariant relative orbits for spacecraft flying formations, and it seems worth to sketch some of the most relevant works. Schaub and Alfriend [1] presented a method to establish J_2 invariant relative orbits for spacecraft formation flying applications. They designed relative orbit geometry using differences in mean orbit elements. Two constraints on the three momenta element differences are derived. Zhang and Dai [2] removed the drifts by adjusting the semi-axis of the follower satellite and obtained a similar conclusion. By means of Routh transformation and dynamical system theory, Koon and Marsden [3] developed a method to find the J_2 invariant orbit. Then Li and Li [4] and Meng *et al.* [5] concluded, from the point of view of relative orbital elements, that the drifts of relative orbit result from the orbital inclination and right ascension of ascending node of the two satellites. Biggs and Becerra [6] proposed a me-

thod to determinate the J_2 invariant orbit with the leader's orbit of zero inclination based on the targeting method in chaos dynamics. Abd El-Salam *et al.* [7] used the Hamiltonian framework to construct an analytical method to design invariant relative constellation orbits due to the zonal harmonics J_2 ; J_3 ; J_4 up to the second order, assuming J_2 being of order 1.

Our propose was to extend Schaub and Alfriend [1] and Abd El-Salam *et al.* [7] model by adding the effect of the third body which have important at high altitude. Using the Hamiltonian framework, the perturbations can be easily added. The Hamiltonian of the problem was constructed by considering the effect of the third body of $\mathcal{O}(J_2^2)$. The expressions for the time rate of change of the secular elements are obtained, second order conditions are established between the differences in momenta elements (semi-major axis, eccentricity and inclination angle) that guarantee that the drift rates of two neighboring orbits are equal on the average.

2. Hamiltonian Approach

There are several ways to derive the equations of motion for any such system. We emphasized on the Hamiltonian structure for this system. The Hamiltonian formulation allows for additional conservative forces to be added to the Hamiltonian, thus the addition of complexity to the model can be incorporated with ease. Non-conservative forces can be added in the momenta equations of motion. The Hamiltonian equations of motion allows us to directly

use control and simulation techniques.

Notations in the whole text, we use the well-known keplerian elements: the semi-major axis a , the eccentricity e , the inclination i , the right ascension of ascending node Ω , the argument of perigee ω , and the mean anomaly M . We also use the true anomaly f and an intermediary variable $\eta = \sqrt{1-e^2}$.

The Hamiltonian in the present framework can be written in the form

$$\mathcal{H} = \frac{1}{2} p^2 - U_{\oplus} + U_t \tag{1}$$

where U_{\oplus} is the force function due to the Earth's gravitational potential, and p is the canonical momentum

$$U_{\oplus} = -\frac{\mu}{r} \sum_{n=0}^{\infty} \sum_{m=0}^{\infty} \left(\frac{R_e}{r}\right)^n P_n^m(\sin \delta) [C_{nm} \cos m\lambda + S_{nm} \sin m\lambda]$$

where R_e is the equatorial radius of the Earth,

$\mu = g'm_e$ is the Earth's gravitational parameter where g' is the gravitational constant;

(r, λ, δ) are the geocentric coordinates of the satellite with λ measured east of Greenwich;

C_{nm} and S_{nm} are harmonic coefficients;

vector and U_t the disturbing function due to the effect of perturbing body.

2.1. Influence of Oblateness Perturbations

The actual shape of the Earth is that of an eggplant. The center of mass does not lie on the spin axis and neither the meridian nor the latitudinal contours are circles. The net result of this irregular shape is to produce a variation in the gravitational acceleration to that predicted using a point mass distribution. The Earth's gravitational potential is usually expressed by the following expression (Vinti's potential)

$P_n^m(\sin \delta)$ are associated Legendre Polynomials.

In the potential function, the terms with $m=0$, $0 < m < n$ and $m=n$ correspond respectively to zonal, tesseral and sectorial harmonics. The Earth gravitational potential can be rewritten, up to second order in J_2 , truncating the series at $n=4$, as, Abd El-Salam *et al.* [7]

$$U_{\oplus} = -\frac{\mu}{r} \left\{ 1 - \frac{1}{4} J_2 \left(\frac{R_e}{r}\right)^2 \left[(3S^2 - 2) - 3S^2 \cos 2(f + \omega) \right] - \frac{1}{8} J_3 \left(\frac{R_e}{r}\right)^3 \left[(15S^3 - 12S) \sin(f + \omega) - 5S^3 \sin 3(f + \omega) \right] - \frac{1}{64} J_4 \left(\frac{R_e}{r}\right)^4 \left[(24 - 120S^2 + 105S^4) + (120S^2 - 140S^4) \cos 4(f + \omega) \right] \right\} \tag{2}$$

where $S = \sin i$ and J_2, J_3, J_4 is the zonal harmonic coefficients.

2.2. Third Body Perturbation

The effect of the third body in the motion of an artificial satellite have become particularly interesting now, when space debris imposes a serious threat to space activities. These perturbations are the most important mechanism of delivering major Earth orbiting objects into the regions where the atmosphere can start their decay.

If it is assumed that the main body; Earth; with mass m_e is fixed in the center of the reference system x-y. The perturbing body, with mass m' is in an elliptic orbit with semi-major axis, a' , eccentricity e' , and mean mo-

tion n' , given by the expression $n'^2 a'^3 = g'(m_e + m')$, r and r' are the radius vectors of the satellite and m' (assuming $r' \gg r_0$), and Ψ is the angle between these radius vectors. The disturbing function (using the tradition expansion in Legendre polynomials) due to the third body is given by, Domingos *et al.* [8],

$$U_t = \frac{\mu' n'^2 a'^2}{2} \left(\frac{a'}{r'}\right)^3 \left(\frac{r}{a}\right)^2 \left[3 \cos^2(\Psi) - 1 \right] \tag{3}$$

where $\mu' = \frac{m'}{m_e + m'}$ and

$$\cos(\Psi) = \alpha \cos(f) + \beta \sin(f)$$

with

$$\alpha = \cos \omega \cos(\Omega - f' - \omega') - \cos i \sin \omega \sin(\Omega - f' - \omega')$$

$$\beta = -\sin \omega \cos(\Omega - f' - \omega') - \cos i \cos \omega \sin(\Omega - f' - \omega')$$

Using the Delaunay canonical-variables (l, g, h, L, G, H) defined by

$$l = M, \text{ Mean anomaly } L = \sqrt{\mu a}$$

$$g = \omega, \text{ Argument of the Perigee } G = L\sqrt{1-e^2} = L\eta$$

$$h = \Omega, \text{ Longitude of ascending node } H = G \cos i$$

Considering J_2 as a small parameter of the problem, the orders of magnitude, up to the second order, of the involved parameters are defined as follows:

$J_2; n' \equiv \mathcal{O}(1), J_3; J_4 \equiv \mathcal{O}(2)$, and let us define the dimensionless parameters as

$$A_{12} = \mu^4 R_e^2, A_{23} = \frac{2\mu^5 R_e^3 J_3}{J_2^2}, A_{24} = \frac{2\mu^6 R_e^4 J_4}{J_2^2}, K = \frac{\mu' n'^2}{J_2^2}$$

The Hamiltonian, Equation (1) up to the second order, can now be expressed as a power series in J_2 as follows

$$\mathcal{H} = \mathcal{H}_0 + \sum_{n=1}^2 \frac{J_2^n}{n!} \mathcal{H}_n \tag{4}$$

where \mathcal{H}_0 represents the unperturbed part of the problem, \mathcal{H}_n is the perturbation:

$$\mathcal{H}_0 = -\frac{\mu^2}{2L^2},$$

$$\mathcal{H}_1 = \frac{A_{12}}{4L^6} \left(\frac{a}{r}\right)^3 \left[(3S^2 - 2) - 3S^2 \cos 2(f + \omega) \right],$$

$$\begin{aligned} \mathcal{H}_2 = & \frac{A_{23}}{8L^8} \left(\frac{a}{r}\right)^4 \left[(15S^3 - 12S) \sin(f + \omega) - 5S^3 \sin 3(f + \omega) \right] \\ & + \frac{A_{24}}{64L^{10}} \left(\frac{a}{r}\right)^5 \left[(24 - 120S^2 + 105S^4) + (120S^2 - 140S^4) \cos 2(f + \omega) + 35S^4 \cos 4(f + \omega) \right] \\ & + Ka^2 \left(\frac{a'}{r'}\right)^3 \left(\frac{r}{a}\right)^2 \left[3(\alpha^2 \cos^2(f) + 2\alpha\beta \cos(f) \sin(f) + \beta^2 \sin^2(f)) - 1 \right] \end{aligned}$$

Now we need to eliminate the short as well as the long periodic terms of the satellite motion in addition to the short periodic terms of the distance perturbing body. Using the perturbation technique based on Lie series and Lie transform, Kamel [9], the transformed Hamiltonian,

for different orders 0, 1, 2 can be written as, Abd El-Salam *et al.* [7] and Domingos *et al.* [8].

$$\mathcal{H}^{**} = \mathcal{H}_0^{**} + \sum_{n=1}^2 \frac{J_2^n}{n!} \mathcal{H}_n^{**} \tag{5}$$

with

$$\mathcal{H}_0^{**} = -\frac{\mu^2}{2L^2}$$

$$\mathcal{H}_1^{**} = \frac{A_{12}}{4L^3 G^3} \left[1 - 3 \frac{H^2}{G^2} \right]$$

$$\begin{aligned} \mathcal{H}_2^{**} = & \frac{3A_{12}^2}{64\mu^2} \frac{1}{L^5 G^7} \left[\left(5 - 10 \frac{H^2}{G^2} - 35 \frac{H^4}{G^4} \right) L^2 + \left(-4 + 24 \frac{H^2}{G^2} - 36 \frac{H^4}{G^4} \right) LG + \left(-5 + 18 \frac{H^2}{G^2} - 5 \frac{H^4}{G^4} \right) G^2 \right] \\ & + \frac{3A_{24}}{128} \frac{1}{L^5 G^7} \left[\left(15 - 150 \frac{H^2}{G^2} + 175 \frac{H^4}{G^4} \right) L^2 + \left(-9 + 90 \frac{H^2}{G^2} - 105 \frac{H^4}{G^4} \right) G^2 \right] + kL^4 \left[\left(5 - 3 \frac{G^2}{L^2} \right) \left(3 \frac{H^2}{G^2} - 1 \right) \right] \end{aligned}$$

where

$$k = \frac{K}{8\mu^2} \left[1 + \frac{3}{2} e'^2 + \frac{15}{8} e'^4 \right]$$

$$\dot{\theta} = \dot{l} + \dot{g} = \frac{\partial \mathcal{H}_0^{**}}{\partial L} + \frac{\partial \mathcal{H}_1^{**}}{\partial G} + \sum_{n=1}^2 \frac{J_2^n}{n!} \left(\frac{\partial \mathcal{H}_n^{**}}{\partial L} + \frac{\partial \mathcal{H}_n^{**}}{\partial G} \right) \tag{6}$$

Using the Hamiltonian canonical equations of the motion, to write $\dot{\theta}$, argument of mean latitude (θ) is the sum of the mean anomaly and the argument of perigee (*i.e.* $\theta = l + g$), as

with

$$\frac{\partial \mathcal{H}_0^{**}}{\partial L} = \frac{\mu^2}{L^3}, \frac{\partial \mathcal{H}_0^{**}}{\partial G} = 0,$$

$$\begin{aligned} \frac{\partial \mathcal{H}_4^{**}}{\partial L} &= \frac{-3A_{12}}{4L^4G^3} \left[1 - 3 \frac{H^2}{G^2} \right] = \frac{-3A_{12}}{4L^7\eta^3} [1 - 3 \cos^2 i], \\ \frac{\partial \mathcal{H}_4^{**}}{\partial G} &= \frac{3A_{12}}{4L^3G^4} \left[\frac{5H^2}{G^2} - 1 \right] = \frac{3A_{12}}{4L^7\eta^4} [5 \cos^2 i - 1], \\ \frac{\partial \mathcal{H}_2^{**}}{\partial L} &= \frac{3A_{12}^2}{64\mu^2} \frac{1}{L^6G^7} \left[-3 \left(5 - 10 \frac{H^2}{G^2} - 35 \frac{H^4}{G^4} \right) L^2 - 4 \left(-4 + 24 \frac{H^2}{G^2} - 36 \frac{H^4}{G^4} \right) LG - 5 \left(-5 + 18 \frac{H^2}{G^2} - 5 \frac{H^4}{G^4} \right) G^2 \right] \\ &\quad + \frac{3A_{24}}{128} \frac{1}{L^4G^7} \left[-3 \left(15 - 150 \frac{H^2}{G^2} + 175 \frac{H^4}{G^4} \right) L^2 - 5 \left(-9 + 90 \frac{H^2}{G^2} - 105 \frac{H^4}{G^4} \right) G^2 \right] + kL^3 \left[\left(20 - 6 \frac{G^2}{L^2} \right) \left(3 \frac{H^2}{G^2} - 1 \right) \right] \\ &= \frac{3A_{12}^2}{64\mu^2} \frac{1}{L^{11}\eta^7} \left[-3(5 - 10 \cos^2 i - 35 \cos^4 i) - 4(-4 + 24 \cos^2 i - 36 \cos^4 i)\eta - 5(-5 + 18 \cos^2 i - 5 \cos^4 i)\eta^2 \right] \\ &\quad + \frac{3A_{24}}{128} \frac{1}{L^9\eta^7} \left[-3(15 - 150 \cos^2 i + 175 \cos^4 i) - 5(-9 + 90 \cos^2 i - 105 \cos^4 i)\eta^2 \right] + kL^3 \left[(20 - 6\eta^2)(3 \cos^2 i - 1) \right] \\ \frac{\partial \mathcal{H}_2^{**}}{\partial G} &= \frac{6A_{12}^2}{128\mu^2} \frac{1}{L^5G^8} \left[5 \left(-7 + 18 \frac{H^2}{G^2} + 77 \frac{H^4}{G^4} \right) L^2 + 24 \left(1 - 8 \frac{H^2}{G^2} + 15 \frac{H^4}{G^4} \right) LG + \left(25 - 126 \frac{H^2}{G^2} + 45 \frac{H^4}{G^4} \right) G^2 \right] \\ &\quad + \frac{15A_{24}}{128} \frac{1}{L^8G^5} \left[\left(-21 + 270 \frac{H^2}{G^2} - 385 \frac{H^4}{G^4} \right) L^2 + \left(9 - 126 \frac{H^2}{G^2} + 189 \frac{H^4}{G^4} \right) G^2 \right] + 6k \frac{L^2}{G^3} [G^4 - 5H^2L^2] \\ &= \frac{6A_{12}^2}{128\mu^2} \frac{1}{L^{11}\eta^8} \left[5(-7 + 18 \cos^2 i + 77 \cos^4 i) + 24(1 - 8 \cos^2 i + 15 \cos^4 i)\eta + (25 - 126 \cos^2 i + 45 \cos^4 i)\eta^2 \right] \\ &\quad + \frac{15A_{24}}{128} \frac{1}{L^{11}\eta^5} \left[(-21 + 270 \cos^2 i - 385 \cos^4 i) + (9 - 126 \cos^2 i + 189 \cos^4 i)\eta^2 \right] + 6k \frac{L^3}{\eta} [\eta^2 - 5 \cos^2 i] \end{aligned}$$

and the secular drift rates of the longitude of the ascending node, \dot{h} :

$$\dot{h} = \frac{\partial \mathcal{H}_0^{**}}{\partial H} + \sum_{n=1}^2 \frac{J_2^n}{n!} \left(\frac{\partial \mathcal{H}_n^{**}}{\partial H} \right) \tag{7}$$

with

$$\begin{aligned} \frac{\partial \mathcal{H}_0^{**}}{\partial H} &= 0, \\ \frac{\partial \mathcal{H}_1^{**}}{\partial H} &= \frac{-3}{2} A_{12} \frac{H}{L^3G^5} = \frac{-3}{2} A_{12} \frac{\cos i}{L^7\eta^4}, \\ \frac{\partial \mathcal{H}_2^{**}}{\partial H} &= \frac{-3A_{12}^2}{16\mu^2} \frac{H}{L^5G^9} \left[5 \left(1 + 7 \frac{H^2}{G^2} \right) L^2 - 12 \left(1 - 3 \frac{H^2}{G^2} \right) LG + \left(-9 + 5 \frac{H^2}{G^2} \right) G^2 \right] \\ &\quad - \frac{15A_{24}}{32} \frac{H}{L^5G^{11}} \left[(3G^2 - 5L^2)(-3G^2 + 7H^2) \right] + 6kL^2H \left[-3 + \frac{5L^2}{G^2} \right] \\ &= \frac{-3A_{12}^2}{16\mu^2} \frac{\cos \cos i}{L^{11}\eta^8} \left[5(1 + 7 \cos^2 i)L^2 - 12(1 - 3 \cos^2 i)\eta + (-9 + 5 \cos^2 i)\eta^2 \right] \\ &\quad - \frac{15A_{24}}{32} \frac{\cos \cos i}{L^{11}\eta^8} \left[(3\eta^2 - 5)(-3 + 7 \cos^2 \cos^2 i) \right] + 6kL^3\eta \cos \cos i \left[-3 + \frac{5}{\eta^2} \right] \end{aligned}$$

3. Constraints for Invariant Orbits

In order to prevent two neighboring orbits from drifting apart, the average secular growth needs to be equal. Short

period oscillations can be ignored here since these are only “temporary” deviations. The long period rates appear secular over a few weeks and they are $\mathcal{O}(J_2^2)$.

Since the mean angle quantities l, g and h do not directly contribute to the secular growth, their values can be chosen at will. However, the mean momenta values L, G and H (and therefore implicitly a, e and i) must be carefully chosen to match the secular drift rates. To keep the satellites from drifting apart over time, it would be desirable to match all three rates $(\dot{l}, \dot{g}, \dot{h})$. We impose

the condition that the relative average drift rate of the angle between the radius vectors be zero. This results in

$$\dot{h}_i = \dot{h}_j \quad \forall i \neq j \tag{9}$$

$$\dot{\theta}_i = \dot{\theta}_j \quad \forall i \neq j \tag{10}$$

Now $\dot{\theta}$ and \dot{h} can be rewritten as

$$\dot{\theta} = \frac{\mu^2}{L^3} + \sum_{n=1}^2 \sum_{m=0}^4 \sum_{l=-1}^5 \frac{J_2^n}{n!} \{ \chi_{\theta}^{nml} L^{-4n-3} \eta^{l-4n} + \gamma_{\theta}^{nml} L^3 \eta^l \} \cos(mi) \tag{11}$$

$$\dot{h} = \sum_{n=1}^2 \sum_{m=0}^4 \sum_{l=-1}^5 \frac{J_2^n}{n!} \{ \chi_h^{nml} L^{-4n-3} \eta^{l-4n} + \gamma_h^{nml} L^3 \eta^l \} \cos(mi) \tag{12}$$

where the non-vanishing coefficients $\chi_{\theta}^{nml}, \gamma_{\theta}^{nml}, \chi_h^{nml}$ and γ_h^{nml} are computed in Appendix I.

orbit can be written as a series expansion about the reference orbit element, here it is enough to keep the second order only, as

Let the reference mean orbit elements be denoted with the subscript "0". The drift rate $\dot{\theta}_i$ of a neighboring

$$\begin{aligned} \delta \dot{\theta}_i = \dot{\theta}_i - \dot{\theta}_0 = & \delta L \left. \frac{\partial \dot{\theta}_i}{\partial L} \right|_{x=x_0} + \delta \eta \left. \frac{\partial \dot{\theta}_i}{\partial \eta} \right|_{x=x_0} + \delta i \left. \frac{\partial \dot{\theta}_i}{\partial i} \right|_{x=x_0} \\ & + \frac{1}{2!} \left[(\delta L)^2 \left. \frac{\partial^2 \dot{\theta}_i}{\partial L^2} \right|_{x=x_0} + (\delta \eta)^2 \left. \frac{\partial^2 \dot{\theta}_i}{\partial \eta^2} \right|_{x=x_0} + (\delta i)^2 \left. \frac{\partial^2 \dot{\theta}_i}{\partial i^2} \right|_{x=x_0} + \delta L \delta \eta \left. \frac{\partial^2 \dot{\theta}_i}{\partial L \partial \eta} \right|_{x=x_0} + \delta \eta \delta i \left. \frac{\partial^2 \dot{\theta}_i}{\partial \eta \partial i} \right|_{x=x_0} + \delta i \delta L \left. \frac{\partial^2 \dot{\theta}_i}{\partial i \partial L} \right|_{x=x_0} \right] \end{aligned} \tag{13}$$

$$\begin{aligned} \delta \dot{h}_i = \dot{h}_i - \dot{h}_0 = & \delta L \left. \frac{\partial \dot{h}_i}{\partial L} \right|_{x=x_0} + \delta \eta \left. \frac{\partial \dot{h}_i}{\partial \eta} \right|_{x=x_0} + \delta i \left. \frac{\partial \dot{h}_i}{\partial i} \right|_{x=x_0} \\ & + \frac{1}{2!} \left[(\delta L)^2 \left. \frac{\partial^2 \dot{h}_i}{\partial L^2} \right|_{x=x_0} + (\delta \eta)^2 \left. \frac{\partial^2 \dot{h}_i}{\partial \eta^2} \right|_{x=x_0} + (\delta i)^2 \left. \frac{\partial^2 \dot{h}_i}{\partial i^2} \right|_{x=x_0} + \delta L \delta \eta \left. \frac{\partial^2 \dot{h}_i}{\partial L \partial \eta} \right|_{x=x_0} + \delta \eta \delta i \left. \frac{\partial^2 \dot{h}_i}{\partial \eta \partial i} \right|_{x=x_0} + \delta i \delta L \left. \frac{\partial^2 \dot{h}_i}{\partial i \partial L} \right|_{x=x_0} \right] \end{aligned} \tag{14}$$

where we make use of the fact that $\dot{\theta} = \dot{\theta}(L, \eta, i)$ and $\dot{h} = \dot{h}(L, \eta, i)$ only, also supposing that $\delta \theta$ is the difference in mean latitude rates,

$$\begin{aligned} \delta L = L_i - L_0, \delta \eta = \eta_i - \eta_0, \delta i = i_i - i_0, x = L, \eta, i \\ \text{and } x_0 = L_0, \eta_0, i_0 \end{aligned}$$

Note that this theory will lead to an analytical second order conditions on the mean orbit elements. To establish a more precise set of orbit elements satisfying Equations

(9) and (10), either $\delta L, \delta \eta$ or δi could be chosen and the remaining two momenta orbit element differences found through a numerical root solving technique. However, the analytical second order conditions provide reasonably accurate solutions to these two constraints equations and provide a wealth of insight into the behavior of Earth potential and third body effect invariant relative orbits.

The required derivatives can be evaluated as

$$\begin{aligned} \dot{\theta}_L = & -3 \frac{\mu^2}{L^4} + \sum_{n=1}^2 \sum_{m=0}^4 \sum_{l=-1}^5 \frac{J_2^n}{n!} \{ (-4n-3) \chi_{\theta}^{nml} L^{-4n-3-1} \eta^{l-4n} + 3 \gamma_{\theta}^{nml} L^2 \eta^l \} \cos(mi) \\ \dot{\theta}_{\eta} = & \sum_{n=1}^2 \sum_{m=0}^4 \sum_{l=-1}^5 \frac{J_2^n}{n!} \{ (l-4n) \chi_{\theta}^{nml} L^{-4n-3} \eta^{l-4n-1} + l \gamma_{\theta}^{nml} L^3 \eta^{l-1} \} \cos(mi) \\ \dot{\theta}_i = & \sum_{n=1}^2 \sum_{m=1}^4 \sum_{l=-1}^5 \frac{-J_2^n}{n!} m \{ \chi_{\theta}^{nml} L^{-4n-3} \eta^{l-4n} + l \gamma_{\theta}^{nml} L^3 \eta^l \} \sin(mi) \\ \dot{\theta}_{LL} = & 12 \frac{\mu^2}{L^5} + \sum_{n=1}^2 \sum_{m=0}^4 \sum_{l=-1}^5 \frac{J_2^n}{n!} \{ (-4n-3)(-4n-3-1) \chi_{\theta}^{nml} L^{-4n-3-2} \eta^{l-4n} + 6 \gamma_{\theta}^{nml} L \eta^l \} \cos(mi) \end{aligned}$$

$$\begin{aligned} \dot{\theta}_{\eta\eta} &= \sum_{n=1}^2 \sum_{m=0}^4 \sum_{t=-1}^5 \frac{J_2^n}{n!} \left\{ (t-4n)(t-4n-1) \chi_{\theta}^{nmt} L^{-4n-3} \eta^{t-4n-2} + t(t-1) \gamma_{\theta}^{nmt} L^3 \eta^{t-2} \right\} \cos(mi) \\ \dot{\theta}_{ii} &= \sum_{n=1}^2 \sum_{m=0}^4 \sum_{t=-1}^5 \frac{-J_2^n}{n!} m^2 \left\{ \chi_{\theta}^{nmt} L^{-4n-3} \eta^{t-4n} + t \gamma_{\theta}^{nmt} L^3 \eta^t \right\} \cos \cos(mi) \\ \dot{\theta}_{L\eta} &= \sum_{n=1}^2 \sum_{m=0}^4 \sum_{t=-1}^5 \frac{J_2^n}{n!} \left\{ (-4n-3)(t-4n) \chi_{\theta}^{nmt} L^{-4n-3-1} \eta^{t-4n-1} + 3t \gamma_{\theta}^{nmt} L^2 \eta^{t-1} \right\} \cos(mi) \\ \dot{\theta}_{\eta i} &= \sum_{n=1}^2 \sum_{m=1}^4 \sum_{t=-1}^5 \frac{-J_2^n}{n!} m \left\{ (t-4n) \chi_{\theta}^{nmt} L^{-4n-3} \eta^{t-4n-1} + t \gamma_{\theta}^{nmt} L^3 \eta^{t-1} \right\} \sin(mi) \\ \dot{\theta}_{iL} &= \sum_{n=1}^2 \sum_{m=1}^4 \sum_{t=-1}^5 \frac{-J_2^n}{n!} m \left\{ (-4n-3) \chi_{\theta}^{nmt} L^{-4n-3-1} \eta^{t-4n} + 3 \gamma_{\theta}^{nmt} L^2 \eta^t \right\} \sin(mi) \end{aligned}$$

and

$$\begin{aligned} \dot{h}_L &= \sum_{n=1}^2 \sum_{m=0}^4 \sum_{t=-1}^5 \frac{J_2^n}{n!} \left\{ (-4n-3) \chi_h^{nmt} L^{-4n-3-1} \eta^{t-4n} + 3 \gamma_h^{nmt} L^2 \eta^t \right\} \cos(mi) \\ \dot{h}_{\eta} &= \sum_{n=1}^2 \sum_{m=0}^4 \sum_{t=-1}^5 \frac{J_2^n}{n!} \left\{ (t-4n) \chi_h^{nmt} L^{-4n-3} \eta^{t-4n-1} + t \gamma_h^{nmt} L^3 \eta^{t-1} \right\} \cos(mi) \\ \dot{h}_i &= \sum_{n=1}^2 \sum_{m=1}^4 \sum_{t=-1}^5 \frac{-J_2^n}{n!} m \left\{ \chi_h^{nmt} L^{-4n-3} \eta^{t-4n} + t \gamma_h^{nmt} L^3 \eta^t \right\} \sin(mi) \\ \dot{h}_{iL} &= \sum_{n=1}^2 \sum_{m=0}^4 \sum_{t=-1}^5 \frac{J_2^n}{n!} \left\{ (-4n-3)(-4n-3-1) \chi_h^{nmt} L^{-4n-3-2} \eta^{t-4n} + 6 \gamma_h^{nmt} L \eta^t \right\} \cos(mi) \\ \dot{h}_{\eta\eta} &= \sum_{n=1}^2 \sum_{m=0}^4 \sum_{t=-1}^5 \frac{J_2^n}{n!} \left\{ (t-4n)(t-4n-1) \chi_h^{nmt} L^{-4n-3} \eta^{t-4n-2} + t(t-1) \gamma_h^{nmt} L^3 \eta^{t-2} \right\} \cos(mi) \\ \dot{h}_{ii} &= \sum_{n=1}^2 \sum_{m=1}^4 \sum_{t=-1}^5 \frac{-J_2^n}{n!} m^2 \left\{ \chi_h^{nmt} L^{-4n-3} \eta^{t-4n} + t \gamma_h^{nmt} L^3 \eta^t \right\} \cos(mi) \\ \dot{h}_{L\eta} &= \sum_{n=1}^2 \sum_{m=0}^4 \sum_{t=-1}^5 \frac{J_2^n}{n!} \left\{ (-4n-3)(t-4n) \chi_h^{nmt} L^{-4n-3-1} \eta^{t-4n-1} + 3t \gamma_h^{nmt} L^2 \eta^{t-1} \right\} \cos(mi) \\ \dot{h}_{\eta i} &= \sum_{n=1}^2 \sum_{m=1}^4 \sum_{t=-1}^5 \frac{-J_2^n}{n!} m \left\{ (t-4n) \chi_h^{nmt} L^{-4n-3} \eta^{t-4n-1} + t \gamma_h^{nmt} L^3 \eta^{t-1} \right\} \sin(mi) \\ \dot{h}_{iL} &= \sum_{n=1}^2 \sum_{m=1}^4 \sum_{t=-1}^5 \frac{-J_2^n}{n!} m \left\{ (-4n-3) \chi_h^{nmt} L^{-4n-3-1} \eta^{t-4n} + 3 \gamma_h^{nmt} L^2 \eta^t \right\} \sin(mi) \end{aligned}$$

where $\dot{\theta}_x = \frac{\partial \dot{\theta}}{\partial x}$ and $\dot{\theta}_{xx} = \frac{\partial^2 \dot{\theta}}{\partial x \partial x}$ with $x = L, \eta, i$.

To enforce equal drift rates $\dot{\theta}_i$ and \dot{h}_i between neighboring orbits, we must set $\delta \dot{\theta}$ and $\delta \dot{h}$ equal to zero in expanded Equations (13) and (14), yields

$$0 = \frac{1}{2} \dot{\theta}_{LL} (\delta L)^2 + \frac{1}{2} \dot{\theta}_{L\eta} \delta L \delta \eta + \frac{1}{2} \dot{\theta}_{\eta\eta} (\delta \eta)^2 + \left(\dot{\theta}_L + \frac{1}{2} \delta i \dot{\theta}_{iL} \right) \delta L + \left(\dot{\theta}_{\eta} + \frac{1}{2} \delta i \dot{\theta}_{\eta i} \right) \delta \eta + \left(\dot{\theta}_i + \frac{1}{2} \delta i \dot{\theta}_{ii} \right) \delta i \quad (15)$$

$$0 = \frac{1}{2} \dot{h}_{LL} (\delta L)^2 + \frac{1}{2} \dot{h}_{L\eta} \delta L \delta \eta + \frac{1}{2} \dot{h}_{\eta\eta} (\delta \eta)^2 + \left(\dot{h}_L + \frac{1}{2} \delta i \dot{h}_{iL} \right) \delta L + \left(\dot{h}_{\eta} + \frac{1}{2} \delta i \dot{h}_{\eta i} \right) \delta \eta + \left(\dot{h}_i + \frac{1}{2} \delta i \dot{h}_{ii} \right) \delta i \quad (16)$$

Equations (15) and (16) are two simultaneous nonlinear algebraic equations in three unknowns, namely $\delta L, \delta \eta, \delta i$. When one of these three unknowns is as-

sumed known (say δi), these two equations can be solved as:

Multiplying Equation (15) by $1/2 \dot{h}_{LL}$ and Equation

(16) by $1/2\dot{\theta}_{LL}$ and then subtracting yields

$$\delta L = -\frac{c_2(\delta\eta)^2 + c_4\delta\eta + c_5}{c_1\delta\eta + c_3} \tag{17}$$

where

$$\begin{aligned} c_1 &= \frac{1}{2}(\dot{\theta}_{L\eta}\dot{h}_{LL} - \dot{\theta}_{LL}\dot{h}_{L\eta}), \\ c_2 &= \frac{1}{2}(\dot{\theta}_{\eta\eta}\dot{h}_{LL} - \dot{\theta}_{LL}\dot{h}_{\eta\eta}), \\ c_3 &= \left(\dot{\theta}_L + \frac{1}{2}\dot{\theta}_{iL}\delta i\right)\dot{h}_{LL} - \left(\dot{h}_L + \frac{1}{2}\dot{h}_{iL}\delta i\right)\dot{\theta}_{LL}, \\ c_4 &= \left(\dot{\theta}_\eta + \frac{1}{2}\dot{\theta}_{\eta i}\delta i\right)\dot{h}_{LL} - \left(\dot{h}_\eta + \frac{1}{2}\dot{h}_{\eta i}\delta i\right)\dot{\theta}_{LL}, \\ c_5 &= \left[\left(\dot{\theta}_i + \frac{1}{2}\dot{\theta}_{ii}\delta i\right)\dot{h}_{LL} - \left(\dot{h}_i + \frac{1}{2}\dot{h}_{ii}\delta i\right)\dot{\theta}_{LL}\right]\delta i. \end{aligned}$$

Substituting Equation (17) into Equation (15) yields an algebraic equation of fourth degree in $\delta\eta$ only in the form

$$d_1(\delta\eta)^4 + d_2(\delta\eta)^3 + d_3(\delta\eta)^2 + d_4\delta\eta + d_5 = 0 \tag{18}$$

where

$$\begin{aligned} d_1 &= c_2^2\dot{\theta}_{LL} - c_1c_2\dot{\theta}_{L\eta} + c_1^2\dot{\theta}_{\eta\eta}, \\ d_2 &= 2c_2c_4\dot{\theta}_{LL} - (c_2c_3 + c_4c_1)\dot{\theta}_{L\eta} \\ &\quad - c_1c_2\left(\dot{\theta}_L + \frac{1}{2}\dot{\theta}_{iL}\delta i\right) + c_1^2\left(\dot{\theta}_\eta + \frac{1}{2}\dot{\theta}_{\eta i}\delta i\right) + 2c_1c_3\dot{\theta}_{\eta\eta}, \\ d_3 &= (c_4^2 + 2c_2c_5)\dot{\theta}_{LL} - (c_4c_3 + c_5c_1)\dot{\theta}_{L\eta} \\ &\quad + c_3^2\dot{\theta}_{\eta\eta} - (c_2c_3 + c_4c_1)\left(\dot{\theta}_L + \frac{1}{2}\dot{\theta}_{iL}\delta i\right) \\ &\quad + 2c_1c_3\left(\dot{\theta}_\eta + \frac{1}{2}\dot{\theta}_{\eta i}\delta i\right) + c_1^2\left(\dot{\theta}_i + \frac{1}{2}\dot{\theta}_{ii}\delta i\right)\delta i \end{aligned}$$

$$\begin{aligned} (\delta L)_{1,2} &= -\frac{c_2(a^2 + b^2 + c^2 - 2ab \pm 2bc \mp 2ac) + c_4(a - b - c) + c_5}{c_1(a - b - c) + c_3} \\ (\delta L)_{3,4} &= -\frac{c_2(a^2 + b^2 + d^2 + 2ab \mp 2bd \mp 2ad) + c_4(a + b - d) + c_5}{c_1(a + b - d) + c_3} \end{aligned}$$

5. Conclusion

Accurate modeling of relative motion dynamics for initial conditions close to the leader satellite is essential for flying formation. Therefore, the solutions of interest are restricted to a specific set of initial conditions that lead to periodic motion, such that the satellites do not drift apart. This paper showed an analytical expression to secular drift rates due to oblate Earth model, truncating its poten-

$$\begin{aligned} d_4 &= 2c_4c_5\dot{\theta}_{LL} - c_5c_3\dot{\theta}_{L\eta} - (c_4c_3 + c_5c_1)\left(\dot{\theta}_L + \frac{1}{2}\dot{\theta}_{iL}\delta i\right) \\ &\quad + c_3^2\left(\dot{\theta}_\eta + \frac{1}{2}\dot{\theta}_{\eta i}\delta i\right) + 2c_1c_3\left(\dot{\theta}_i + \frac{1}{2}\dot{\theta}_{ii}\delta i\right)\delta i, \\ d_5 &= c_5^2\dot{\theta}_{LL} - c_5c_3\left(\dot{\theta}_L + \frac{1}{2}\dot{\theta}_{iL}\delta i\right) + c_3^2\left(\dot{\theta}_i + \frac{1}{2}\dot{\theta}_{ii}\delta i\right)\delta i. \end{aligned}$$

4. Solution of the Quartic Equation (18)

The roots of the quartic Equation (18) can be written as

$$(\delta\eta)_{1,2} = a + b \pm c, (\delta\eta)_{3,4} = a + b \pm d$$

where

$$\begin{aligned} a &= -\frac{d_2}{4d_1}, b = \frac{1}{2}\sqrt{2e - f - g - h}, \\ c &= \frac{1}{2}\sqrt{e + f + g}, d = \frac{1}{2}\sqrt{2e - f + g + h} \end{aligned}$$

with

$$e = 4a^2 - \frac{2d_3}{3d_1}, f = \frac{\sqrt[3]{i}}{3d_1j}, g = \frac{k}{3*\sqrt[3]{2d_1}}, h = \frac{l}{8b}$$

where

$$\begin{aligned} i &= d_3^2 - 3d_2d_4 + 12d_1d_5, j = \left(m + \sqrt{-4i^3 + m^2}\right), \\ l &= -\left(\frac{d_2}{d_1}\right)^3 + \frac{4d_2d_3}{d_1^2} - \frac{4d_4}{d_1} \end{aligned}$$

and

$$m = 2d_3^3 - 9d_2d_3d_4 + 27d_1d_4^3 + 27d_2^3d_5 - 72d_1d_3d_5.$$

Substituting the four roots $\delta\eta$'s into Equation (17) yields the four constraints δL 's that guarantee the invariance of the relative motion of certain satellite constellation

tial series at J_4 , and third body effect and set it equal between two neighboring orbits. It followed the same steps used before in Abd El-Salam *et al.* [7] for the Earth model so the calculation of Abd El-Salam *et al.* [7] and Schaub and Alfriend [1] is a special case from this calculations. The variation in the inclination (δi) can be chosen at will for the nominal inclination, and the variations in both the eccentricity (δe) and semi-major axis

(δa) from their nominal values are set to zero. Noted that these constraint conditions are not justified near the critical inclination angle. Using η instead of e to avoid the singularity when $e \rightarrow 0$ but for $e = 0$ the nonsingular elements must be used. Future developments of this approach to the formation flying problem include another perturbation forces like solar radiation and luni-solar effects.

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Appendix I

$$\begin{aligned} \chi_{\theta}^{100} &= \frac{9A_{12}}{8}, \chi_{\theta}^{101} = \frac{5}{3}\chi_{\theta}^{100}, \chi_{\theta}^{110} = \frac{1}{3}\chi_{\theta}^{100}, \chi_{\theta}^{111} = \chi_{\theta}^{100} \\ \chi_{\theta}^{200} &= \frac{3705A_{12}^2}{512\mu^2} - \frac{3645A_{24}}{1024}, \chi_{\theta}^{220} = \frac{1425A_{12}^2}{128\mu^2} - \frac{1725A_{24}}{256}, \chi_{\theta}^{240} = \frac{1155A_{12}^2}{512\mu^2} - \frac{5775A_{24}}{1024} \\ \chi_{\theta}^{201} &= \frac{2457A_{12}^2}{512\mu^2} - \frac{405A_{24}}{1024}, \chi_{\theta}^{221} = \frac{909A_{12}^2}{128\mu^2} - \frac{225A_{24}}{256}, \chi_{\theta}^{241} = \frac{1395A_{12}^2}{512\mu^2} - \frac{1575A_{24}}{1024} \\ \chi_{\theta}^{202} &= \frac{21A_{12}^2}{512\mu^2} + \frac{2025A_{24}}{1024}, \chi_{\theta}^{222} = \frac{-99A_{12}^2}{128\mu^2} + \frac{945A_{24}}{256}, \chi_{\theta}^{242} = \frac{567A_{12}^2}{512\mu^2} + \frac{2835A_{24}}{1024} \\ \chi_{\theta}^{203} &= \frac{-255A_{12}^2}{512\mu^2} + \frac{405A_{24}}{1024}, \chi_{\theta}^{220} = \frac{-195A_{12}^2}{128\mu^2} + \frac{225A_{24}}{256}, \chi_{\theta}^{240} = \frac{75A_{12}^2}{512\mu^2} + \frac{1575A_{24}}{1024} \\ \gamma_{\theta}^{20-1} &= -15k, \gamma_{\theta}^{24-1} = \chi_{\theta}^{20-1}, \gamma_{\theta}^{200} = 10k, \gamma_{\theta}^{240} = 3\chi_{\theta}^{200}, \gamma_{\theta}^{201} = 6k, \gamma_{\theta}^{202} = -3k, \gamma_{\theta}^{242} = -9k \\ \chi_h^{110} &= \frac{-3A_{12}}{2}, \chi_h^{210} = \frac{-375A_{12}^2}{64\mu^2} + \frac{675A_{24}}{128}, \chi_h^{230} = \frac{-105A_{12}^2}{64\mu^2} + \frac{525A_{24}}{128}, \chi_h^{211} = \frac{-45A_{12}^2}{16\mu^2}, \chi_h^{231} = \frac{-27A_{12}^2}{16\mu^2} \\ \chi_h^{212} &= \frac{63A_{12}^2}{64\mu^2} - \frac{405A_{24}}{128}, \chi_h^{232} = \frac{-15A_{12}^2}{64\mu^2} - \frac{315A_{24}}{128} \\ \gamma_h^{21-1} &= 30k, \gamma_h^{211} = -18k \end{aligned}$$