



HAL
open science

CdWO scintillating bolometer for Double Beta Decay: Light and Heat anticorrelation, light yield and quenching factors

C. Arnaboldi, J.W. Beeman, O. Cremonesi, L. Gironi, M. Pavan, G. Pessina,
S. Pirro, E. Previtali

► **To cite this version:**

C. Arnaboldi, J.W. Beeman, O. Cremonesi, L. Gironi, M. Pavan, et al.. CdWO scintillating bolometer for Double Beta Decay: Light and Heat anticorrelation, light yield and quenching factors. *Astroparticle Physics*, 2010, 34 (3), pp.143. 10.1016/j.astropartphys.2010.06.009 . hal-00671850

HAL Id: hal-00671850

<https://hal.science/hal-00671850>

Submitted on 19 Feb 2012

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Accepted Manuscript

CdWO₄ scintillating bolometer for Double Beta Decay: Light and Heat anticorrelation, light yield and quenching factors

C. Arnaboldi, J.W. Beeman, O. Cremonesi, L. Gironi, M. Pavan, G. Pessina, S. Pirro, E. Previtali

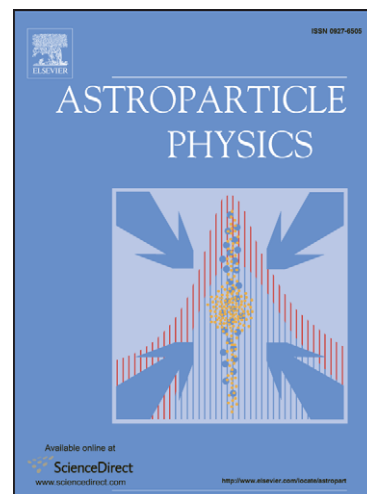
PII: S0927-6505(10)00121-0
DOI: [10.1016/j.astropartphys.2010.06.009](https://doi.org/10.1016/j.astropartphys.2010.06.009)
Reference: ASTPHY 1505

To appear in: *Astroparticle Physics*

Received Date: 8 May 2010
Revised Date: 17 June 2010
Accepted Date: 28 June 2010

Please cite this article as: C. Arnaboldi, J.W. Beeman, O. Cremonesi, L. Gironi, M. Pavan, G. Pessina, S. Pirro, E. Previtali, CdWO₄ scintillating bolometer for Double Beta Decay: Light and Heat anticorrelation, light yield and quenching factors, *Astroparticle Physics* (2010), doi: [10.1016/j.astropartphys.2010.06.009](https://doi.org/10.1016/j.astropartphys.2010.06.009)

This is a PDF file of an unedited manuscript that has been accepted for publication. As a service to our customers we are providing this early version of the manuscript. The manuscript will undergo copyediting, typesetting, and review of the resulting proof before it is published in its final form. Please note that during the production process errors may be discovered which could affect the content, and all legal disclaimers that apply to the journal pertain.



CdWO₄ scintillating bolometer for Double Beta Decay: Light and Heat anticorrelation, light yield and quenching factors

C. Arnaboldi^a, J.W. Beeman^c, O. Cremonesi^a, L. Gironi^{a,b}, M. Pavan^{a,b}, G. Pessina^a, S. Pirro^{*a}, E. Previtali^a

^aINFN - Sezione di Milano Bicocca I 20126 Milano - Italy

^bDipartimento di Fisica - Università di Milano Bicocca I 20126 Milano - Italy

^cLawrence Berkeley National Laboratory, Berkeley, California 94720, USA

Abstract

We report the performances of a 0.51 kg CdWO₄ scintillating bolometer to be used for future Double Beta Decay Experiments. The simultaneous read-out of the heat and the scintillation light allows to discriminate between different interacting particles aiming at the disentanglement and the reduction of background contribution, key issue for next generation experiments. We will describe the observed anticorrelation between the heat and the light signal and we will show how this feature can be used in order to increase the energy resolution of the bolometer over the entire energy spectrum, improving up to a factor 2.6 on the 2615 keV line of ²⁰⁸Tl. The detector was tested in a 433 h background measurement that permitted to estimate extremely low internal trace contaminations of ²³²Th and ²³⁸U. The light yield of γ/β , α 's and neutrons is presented. Furthermore we developed a method in order to correctly evaluate the absolute thermal quenching factor of α -particles in scintillating bolometers.

Key words:

Double Beta Decay, Bolometers, CdWO₄, Quenching Factor

1. Introduction

Double Beta Decay (DBD) searches became of critical importance after the discovery of the neutrino oscillations and plenty of experiments are now in the construction phase and many others are in R&D phase [1, 2, 3, 4]. The main challenges for all the different experimental techniques are the same [5]: i) increase the active mass, ii) decrease the background, and iii) increase the energy resolution.

Thermal bolometers are ideal detectors for this survey: crystals can be grown with different interesting DBD-emitters and, fundamental for next generation experiments, they show an excellent energy resolution.

The CUORICINO Experiment [6], constituted by an array of 62 TeO₂ crystal bolometers, demonstrated not only the power of this technique but also that the main source of background for these detectors arises from surface contaminations of radioactive α -emitters. Moreover simulations show that this contribution will largely dominate the expected background of the CUORE Experiment [7, 8] in the region of interest, since there is no possibility to separate this background from the two DBD electrons. The natural way to discriminate this background is to use a scintillating bolometer [9]. In such a device the simultaneous and independent read out of the heat and the scintillation light permits to discriminate events due to γ/β , α and neutrons thanks to their different scintillation yield. Moreover if the crystal is based on a DBD emitter whose transition energy exceeds

the 2615 keV γ -line of ²⁰⁸Tl then the environmental background due to natural γ 's will decrease abruptly.

CdWO₄ is an ideal candidate for such kind of detector:

- it is a well established scintillator
- ¹¹⁶Cd (7.5 % i.a) has a DBD transition at 2805 keV
- the light yield (LY) is rather large
- the radiopurity of this compound is "naturally" high

Due to these favourable features this crystal compound was already used to perform a DBD experiment [10] using "standard" Photomultipliers. A large mass experiment, based on enriched CdWO₄ crystals, readout by Photomultipliers, was also proposed [11]. But this technique, limited by the modest achievable energy resolution, key point of future experiments, is no more pursued.

¹¹⁶Cd is also used in different DBD Experiments. The Nemo Experiment [12] uses 405 g of enriched ¹¹⁶Cd thin foils in a calorimetric-tracker detector (mainly for the study of the 2 neutrino decay spectrum). The COBRA collaboration is developing small CdZnTe semiconductor detectors [13] for a future DBD Experiment.

2. Experimental details

The dimensions of the CdWO₄ crystal tested are 4 cm diameter, 5 cm height. All the surfaces of the crystal are polished at optical grade. The crystal was tested as a standard scintillator at

*Corresponding author

Email address: Stefano.Pirro@mib.infn.it (S. Pirro)

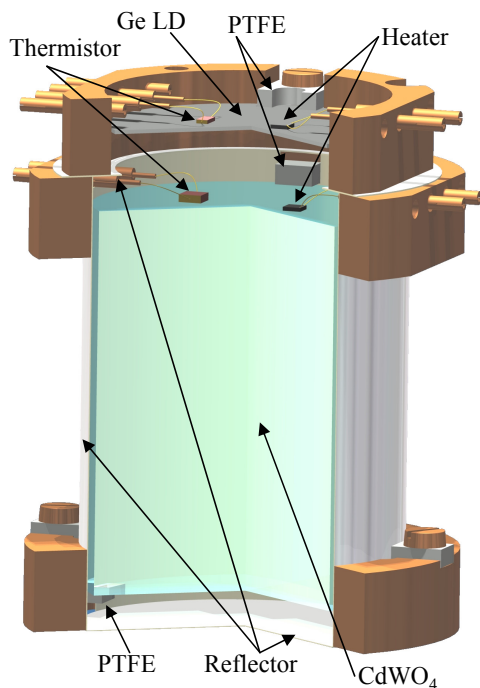


Figure 1: Setup of the detector. The two Cu columns, holding the two frames, are not visible from the chosen perspective.

room temperature. It was wrapped with a reflecting sheet (3M Radiant Mirror film VM2000) and coupled with optical grease to a Photomultiplier (Hamamatsu R6233). The energy resolution evaluated on the ^{137}Cs line is 12.3 % FWHM.

Our detector setup is schematized in Fig. 1. It is held by means of four L-shaped Teflon pieces fixed to the two cylindrical Cu frames. The frames are held together through two Cu columns. The crystal is surrounded (without being in thermal contact) by a 43 mm diameter reflecting sheet (3M VM2002). The Light Detector (LD) [14] is constituted by a 36 mm diameter, 1 mm thick pure Ge crystal absorber. The surface of the Ge facing the crystal is further “darkened” through the deposition of a 600 Å layer of SiO_2 in order to increase the light absorption. Furthermore since the diameter of the LD is smaller with respect to the one of the reflecting cavity, a “reflecting ring” has been mounted in order to decrease light losses. On the opposite face of the crystal a reflecting sheet is mounted. The temperature sensor of the CdWO_4 crystal is a $3 \times 3 \times 1 \text{ mm}^3$ neutron transmutation doped Germanium, the same used in the CUORICINO experiment. The temperature sensor of the LD has smaller volume ($3 \times 1.5 \times 0.4 \text{ mm}^3$) in order to decrease its heat capacity, increasing therefore its thermal signal. A resistor of $\sim 300 \text{ k}\Omega$, realized with a heavily doped meander on a 3.5 mm^3 silicon chip, is attached to each absorber and acts as a heater to stabilize the gain of the bolometer [15, 16]. The detectors were operated deep underground in the Gran Sasso National Laboratories in the CUORE R&D test cryostat. The details of the electronics and the cryogenic facility can be found elsewhere [17, 18, 19].

The heat and light pulses, produced by a particle interacting

in the CdWO_4 crystal and transduced in a voltage pulse by the NTD thermistors, are amplified and fed into a 16 bit NI 6225 USB ADC unit. The entire waveform of each triggered voltage pulse is sampled and acquired. The time window has a width of 256 ms sampled with 512 points. The trigger of the CdWO_4 is software generated while the LD is automatically acquired in coincidence with the former. The amplitude and the shape of the voltage pulse is then determined by the off line analysis that makes use of the Optimal Filter technique. The energy calibration of the CdWO_4 crystal is performed using γ sources placed outside the cryostat. The Heat axis is calibrated attributing to each identified γ peak the nominal energy of the line, as if all the energy is converted into heat. Consequently this calibration does not provide an absolute evaluation of the heat deposited in the crystal. The dependency of amplitude from energy is parameterized with a second order polynomial in $\log(V)$ where V is the heat pulse amplitude. The three coefficients of the polynomial are fitted on calibration data (the heat pulse amplitude corresponding to the known γ lines visible in the spectrum). The choice of such a function was established by means of simulation studies based on a thermal model of the detectors.

The energy calibration of the LD is obtained thanks to a weak ^{55}Fe source placed close to the Ge that illuminates homogeneously the face opposed to the CdWO_4 crystal. During the LD calibration its trigger is set independent from the one of the CdWO_4 . The LD is calibrated using a simple linear function. The FWHM energy resolution of the LD, evaluated on the X doublet at 5.90 and 6.49 keV, is 480 eV FWHM. In the range of 0–50 keV, i.e. the energy interval of the light signals, the response of the LD can be definitely assumed to be linear.

Three sets of data have been collected with this device: two calibrations using ^{232}Th and ^{40}K sources, a long background measurement (433 h) and a neutron measurement (8 h) done exposing the detector to an Am-Be source.

3. Light-Heat scatter plot

In this field the usual way to present the results is to draw the Light vs. Heat scatter plot. Here each event is identified by a point with abscissa equal to the heat signal (recorded by the CdWO_4 bolometer), and ordinate equal to the light signal (contemporarily recorded by the LD). In the scatter plot, γ/β , α and neutrons give rise to separate bands, in virtue of their characteristic Light to Heat ratio (see Fig. 2). This feature is the result of the different LY’s characterizing these particles: γ/β events belong to a distribution (see Fig. 3) characterized by the $\theta_{\beta\gamma}$ angle, while the α events are characterized by θ_α . The scintillation Quenching Factor (QF) is defined as the ratio of the scintillating yield of an interacting particle (α , neutron, nucleus) with respect to the LY of a γ/β event at the same energy. From Fig. 3 we have $\text{QF}_\alpha = \tan(\theta_\alpha)/\tan(\theta_{\beta\gamma})$. But this holds only in first approximation, as will be exposed in Sec. 5 and Sec. 6.

Within each band, monochromatic events (i.e. those corresponding to the complete absorption of a monochromatic particle in the CdWO_4 crystal) appear as sections of lines with negative slope, showing a strong anticorrelation between Heat and

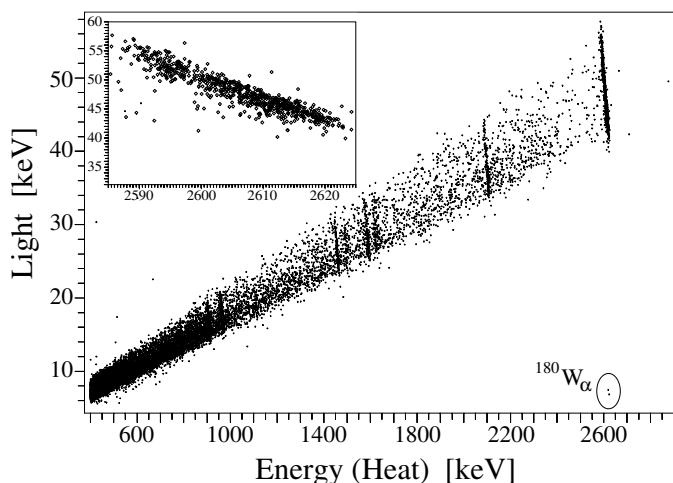


Figure 2: Scatter plot Light vs. Heat obtained in a 96 h calibration using an external ^{232}Th source. In the inset the highlight of the ^{208}Tl line. In the circle two events due to the internal α -decay of ^{180}W . The energy spectrum below 400 keV is completely dominated by the ^{113}Cd β -decay, with a rate close to 0.4 Hz.

Light. This is evident in the Light vs. Heat scatter plot obtained with a source of ^{232}Th (Fig. 2) where different γ -lines are clearly visible. Despite the much lower statistics, the same feature is perceptible also for the α lines that appear in the background spectrum (Fig. 7).

A simple model accounts for the observed pattern. The energy E of a monochromatic particle, fully absorbed in the CdWO_4 crystal, is divided into two channels: a fraction (L) is spent in the production of Light (photons) and a fraction (H) is spent in the production of Heat (phonons). In absence of “blind” channels (i.e. channels in which the energy deposited into the crystal is stored in some system that do not takes part in signal formation) the energy conservation requires:

$$E = E_{\text{Heat}} + E_{\text{Light}} = (1 - k)E + kE = H + L \quad (1)$$

being k the fraction of the total energy that escapes the crystal in form of light. The last equation states not only the energy conservation, but also a subtle distinction between the Heat (what is measured) and the energy definition.

Monochromatic events should produce a spot in the scatter plot, with a size determined by the intrinsic energy resolution of the heat and light detectors. The observed spread along a negative slope line, however, can be accounted for assuming the existence of fluctuations in the H and L signal amplitudes. Since the energy has to be conserved the amplitude of the fluctuations (δL and δH) must compensate each other: $\delta L = -\delta H$. In other words the fluctuations are anti-correlated between each other and produce the observed pattern. They are visible because their amplitude is much larger than the intrinsic resolution of the detectors.

We can devise two mechanisms that introduce a fluctuation in the L/H ratio: i) the statistical Poissonian fluctuation in the emitted light: the more the energy spent in light production, the less the energy spent in heat; ii) any variation of the LY as those due to position dependent effects (inhomogeneities and defects

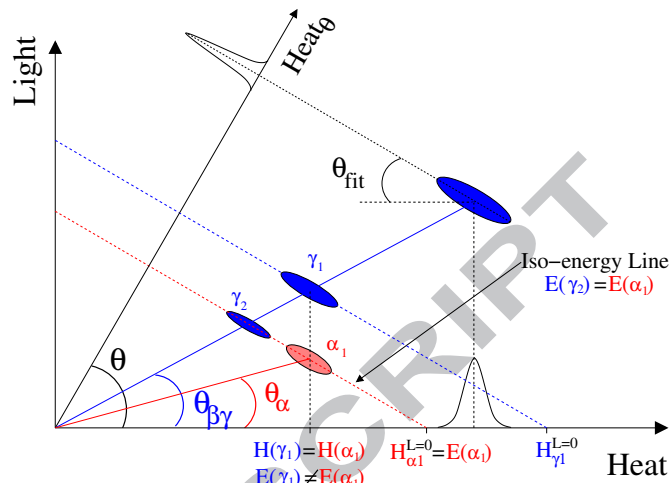


Figure 3: Light vs. Heat scheme for the interpretation of the energy correction method. The blue (dark) spots represent the γ/β monochromatic events, while the red (light) represents an α .

of the crystals that modify the light emission or self absorption of scintillation photons). Regardless the mechanism, we will demonstrate, in the next section, how this anticorrelation can be used in order to increase (correct) the energy resolution of the detector over the entire energy spectrum.

Fig. 3 tries to summarize the model here discussed. The dark spots represent monochromatic energy deposition in the CdWO_4 crystal, as those observed during a γ calibration. For each of them a negative slope (iso-energy) line is tracked. The intercept on the Heat axis ($H^{\text{Light}=0}$) is the Heat that would correspond to a full heat conversion of the deposited energy (in absence of light emission or in the case in which all the light is absorbed by the crystal itself). In this simple model γ/β 's and α 's releasing the same energy within the CdWO_4 crystal lie on the same iso-energy line. But due to the larger LY, γ/β events will convert less heat within the crystal with respect to an α particle releasing the same energy: part of the energy escapes the crystal in form of photons. This feature can be easily deduced from Fig. 3: the α_1 particle release the same total energy of the γ_2 particle (they belong to the same iso-energy line) but they show a different position on the Heat axis. In fact the α_1 shows the same heat of the γ_1 particle that releases a larger energy within (but not “into”) the crystal. The “usual” picture (in which γ/β and α of the same energy show the same heat signal) corresponds to the limit in which the light emitted is negligible (or completely re-absorbed by the crystal): in this case the distinction between H and E becomes meaningless.

It is important to remark that up to now it was assumed to be able to measure the *absolute* values of energy converted into heat and light.

Actually, the measured experimental heat and light signals can be written as

$$\begin{aligned} H &= \alpha(1 - k)E \\ L &= \beta\epsilon kE \end{aligned} \quad (2)$$

being α and β the absolute calibration factors for the Heat and

Light axis. The factor ϵ , instead, takes into account the overall light collection efficiency.

Under the condition that α and β are not depending on the energy, and considering that the conservation of energy requires that $\delta[kE] = -\delta[(1-k)E]$, Eq. 2 implies that

$$\frac{\delta L}{\delta H} = -\frac{\beta\epsilon}{\alpha} = -\tan(\theta_{fit}) \quad (3)$$

The last equation states that the effect of miss-calibrations of the Heat and Light axis simply implies a variation of the slope of monochromatic lines, which is exactly what we observe experimentally. Also the light collection efficiency plays a fundamental role in the evaluation of the slope, as discussed in the next section.

It is important to note that, unlike the Heat channel, the LD energy calibration is more delicate. The energy calibration on this detector is performed using *ionizing* X-Rays from the ^{55}Fe source. *A priori* the thermal signal that arises by the absorption of 1 keV of photons at the boundaries of the Ge crystal (moreover covered by the SiO_2 layer) could give a different thermal signal with respect to a 1 keV ionizing energy. In any case this effect still preserves proportionality, so that it can be considered within the factor β .

4. Light-Heat anticorrelation: energy resolution

The energy anticorrelation between the two signals of a double readout system was already observed in ionization/scintillation detectors [20] as well as in heat/scintillation bolometers [21] with a device very similar to the one presented here. In the former case this anticorrelation was demonstrated to improve the energy resolution by a factor close to $\sim 25\%$ (evaluated @570 keV). In paper [21] two different γ -lines were studied at rather low energy. An evident anticorrelation was found but, however, the correlation factor of the two lines was found to strongly depend on the energy.

The energy correction procedure is graphically explained in Fig. 3. Here blue (dark) and red (light) spots represent γ/β and α produced by monochromatic events. In order to obtain a spectrum with improved energy resolution we have to combine the heat and light values of each single event in an appropriate way. This can be obtained in a natural way performing a simple SU(1) rotation in the scatter plane and projecting the points on the Heat_θ axis:

$$\begin{aligned} \text{Heat}_\theta &= \text{Heat} \cos\theta + \text{Light} \sin\theta \\ \text{Ligh}_\theta &= -\text{Heat} \sin\theta + \text{Light} \cos\theta \end{aligned} \quad (4)$$

The value of θ can be evaluated in two different ways: i) it can be “optimized” in order to obtain the best energy resolution on the Heat_θ axis (energy minimization); ii) by fitting the monochromatic spots in the scatter plot with a negative slope line, evaluating the θ_{fit} angle for each single distribution (for construction we have $\theta = \pi/2 - \theta_{fit}$)

It is clear from this scheme that some assumptions should be satisfied in order for this procedure to work properly:

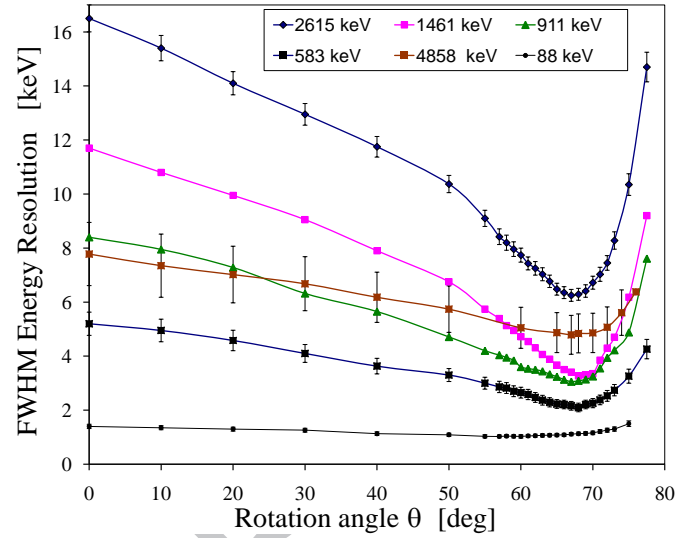


Figure 4: FWHM energy resolutions evaluated for different θ angles. For each angle a spectrum is produced and the corresponding peaks are fitted. The curves through the data points are only to guide eyes. For two peaks the corresponding error on the FWHM is not plotted for better clarity.

- the slope of each monochromatic spot (i.e. the θ_{fit} angle) or the θ angle minimizing the FWHM has to be the same within all the lines;
- the LY of the different class of interacting particles (α 's and γ/β) has to be independent from energy.

The first condition ensures that the rotation minimizes the FWHM on the *whole* energy spectrum. It has to be noted that if the first condition is not verified, this technique still provides a useful tool for resolution optimization. In case of different θ_{fit} values the rotation angle will be chosen in order to reach the best performances in the region of more interest for the physics (e.g. the region where the DBD peak should appear).

The second condition ensures that the projected spectrum keeps linearity and energy calibration. This point is less trivial and will be treated in more detail in Sec. 5.

We applied both the methods in order to evaluate the rotation angle. Using the energy minimization method, the value of the rotation θ angle is allowed to vary over a wide range (0-80 degrees) and for each angle a spectrum is produced projecting the data on the Heat_θ axis. Finally the FWHM energy resolution of the more intense peaks is evaluated by a fit (we use an asymmetric Gaussian reporting then the average FWHM).

We applied this method not only to the calibration spectrum but also to the background measurement. In this way we were able to study also the 88 keV γ -line of ^{109}Cd (accidentally contamination of our CdWO_4 sample) and the α peaks due to internal trace contaminations: in this latter case we observe a line at the full energy of the decay ($Q=\alpha+\text{nuclear recoil}$). The internal α -lines observed are $^{234}\text{U}-\alpha$ ($Q=4858$ keV), $^{180}\text{W}-\alpha$ ($Q=2516$ keV) and $^{238}\text{U}-\alpha$ ($Q=4270$ keV).

In Fig. 4 we report the FWHM energy resolutions evaluated on the most intense lines obtained varying the rotation angle θ

and projecting on Heat_θ . The energy resolutions of the peaks are evaluated for each single rotated spectrum. The minimum of the energy resolution occurs within $\theta=(67.5 \pm 0.5)$ deg for all the lines, except for the 88 keV Electron Capture internal line of ^{109}Cd .

In fact the minimum of the energy resolution for this line, 1.03 ± 0.06 keV FWHM is obtained at $\theta=60$ deg while at 67.5 deg it becomes 1.11 ± 0.06 keV. These two values seem rather similar, but from Fig. 4 it turns out that the resolution curve for this peak shows a different behaviour with respect to all the other lines. We do not have an explanation for this enhancement in the anticorrelation at low energies, but the same effect was already observed in [21].

Each monochromatic events distribution in the scatter plot was also fitted with a linear function, as previously discussed.

The θ_{fit} angles as well as the most relevant parameters of this analysis are summarized in Tab. 1. The values of the θ_{fit} angles for the α 's and γ/β 's lines reported here are compatible within the experimental errors within the two groups (with the exception of the 88 keV line). However, with respect to the γ/β 's, the angular coefficient calculated by fitting linearly the single lines gives a result that is systematically slightly larger with respect to the value (67.5 ± 0.5) obtained through the energy minimization. This systematic slightly difference could arise from the fact that while the bidimensional fit is model independent, the energy minimization depends on the fitting procedure of the peaks. Moreover it should be pointed out that the quoted error on θ_{fit} don't include systematic effects (introduced, for example, by the continuum background) whose values could dominate the total error.

The values of the θ_{fit} angles reported in Tab. 1 seems to show a slightly difference between α and γ/β . However it has to be pointed out that the quoted error on the θ_{fit} for the α particles has a (further) systematic error that cannot be easily evaluated. The point is that the α events were collected during the long background measurement. During this period the CdWO_4 crystal was slowly cooling down; this happen very often with cryogenic detectors and is due to residual heat leaks of the materials constituting the overall setup. The base (stable) temperature is normally reached after 2-3 weeks. In our case the background measurement started immediately after the down cooling so that the initial measurements suffered by this (large) drift (the LD, on the contrary, was rather stable during the same period). This drift, as explained in Sec. 2, is usually corrected through the Heater pulse. If the temperature drift is rather small ($\lesssim 50 \mu\text{K}$) then the correction is independent from the energy. On the contrary, if the drift is larger (in this case the drift was $\approx 0.3 \text{ mK}$) then the shape of the signal changes appreciably. In such cases the correction starts to depend on the energy. As a consequence a small drift in function of the time (or the temperature) becomes appreciable. Moreover this drift depends on the energy. So, in order to evaluate the second, third and last column in Tab. 1 each α peak was "self-stabilized". This can introduce an uncontrolled systematics since we are somehow forcing a distribution to obtain a minimum. On the other hand, a check was made on the weak 2615 keV γ -line present in the background spectrum. As in the case of the α 's, we had to self-correct the

drift. The obtained θ_{fit} value is 71.5 ± 0.4 , rather consistent with the α background. For the following we will assume $\theta_{fit}^{\gamma/\beta} = \theta_{fit}^\alpha$. The ^{232}Th and ^{40}K calibrations, as well as the 88 keV analysis, on the other hand, were made at the end of the measurement in a rather stable working temperature.

As a final remark it has to be noticed that the projection of the events on the Heat_θ axis introduces an Heat (energy) normalization factor K that will be different for different classes of particles, namely $K_{\gamma/\beta,\alpha} = \cos\theta(1 + \tan\theta \tan\theta_{\gamma/\beta,\alpha})$. This feature is rather evident in Fig. 3: in the Heat axis we have $H(\alpha 1) = H(\gamma 1)$ while in the rotated Heat_θ axis we have $H(\alpha 1) = H(\gamma 2)$. This has two consequences: the first, obvious, is that the intercalibration between α and γ/β is different within the Heat spectrum and the Heat_θ spectrum. The second is that the same consideration holds within the same class of events: if the LY of α particles (for example) depends slightly on energy, then the linearity of the projected spectrum will change.

5. Light Yield and Scintillation Quenching Factor

In Tab. 1 we also present the values of the LY, determined on the different monochromatic lines. We define the LY as the energy released in the LD (in keV) for a nominal energy deposition of 1 MeV in the scintillating crystal.

Table 1: Table with the main parameters of the detector. The theoretical FWHM energy resolution of the CdWO_4 crystal, evaluated through Optimal Filtering, is 0.8 keV. The LY of the α -particles (last three lines) is evaluated for the overall Q-value of the decay. The correct value (assuming negligible scintillation for the nucleus recoil) is $\approx 2\%$ larger. †See text.

Line	FWHM $_{\theta=0}$	FWHM $_{\theta=67.5}$	LY	$\pi/2 - \theta_{fit}$
[keV]	[keV]	[keV]	[keV/MeV]	[deg]
88	1.4 ± 0.08	1.11 ± 0.06	17.3 ± 0.2	61.6 ± 0.9
583	5.2 ± 0.4	2.1 ± 0.2	17.0 ± 0.35	70.8 ± 0.4
911	8.4 ± 0.5	3.05 ± 0.2	17.5 ± 0.2	70.6 ± 0.3
1461	11.7 ± 0.5	3.27 ± 0.2	17.45 ± 0.15	70.4 ± 0.3
2615	16.5 ± 0.5	6.25 ± 0.22	17.6 ± 0.1	69.8 ± 0.2
2516	5.2 ± 1.6 †	4.3 ± 1.3 †	2.74 ± 0.08	73.0 ± 0.2 †
4270	7.0 ± 1.5 †	4.4 ± 1.0 †	3.29 ± 0.07	72.3 ± 0.2 †
4858	7.8 ± 1.2 †	4.8 ± 0.7 †	3.44 ± 0.08	72.7 ± 0.2 †

The table shows that the LY of the γ/β events is, within the error, the same for all the observed lines (including the 88 keV line that shows a different resolution curve in Fig. 4). On the other hand a small, but rather clear, energy dependence is evident for α particles. This effect can also be observed in Fig. 6 in which it is evident that the α particles belong to a curve rather than a straight line. Moreover, as already stated, the observed α -lines are internal so that the effect cannot be ascribed to surface effects. A theoretical explanation for this behaviour and a detailed discussion of the energy dependence of LY and QF in different scintillating crystals can be found in [22]. Essentially, the observed behaviour reflects the fact that an α particle scintillates less with respect to an electron because of its larger dE/dx which can induce saturation effects in the scintillator (the

Birks law [23]). The energy dependence of the LY and QF is a consequence of the energy dependence of the stopping power. Electrons, due to their low stopping power, do not suffer of saturation effects and their LY is, consequently, energy independent. For this reason the definition of $QF_\alpha = \tan(\theta_\alpha)/\tan(\theta_{\beta\gamma})$, given in Sec. 3, represents only an approximation.

Using the values in the table we can evaluate the scintillation Quenching Factors of α particles ($QF = LY(\alpha)/LY(\beta)$): $QF_{180W} = 0.160 \pm 0.006$, $QF_{238U} = 0.192 \pm 0.006$, $QF_{234U} = 0.201 \pm 0.006$.

Within this framework, furthermore, it is possible to evaluate the total amount of light k that escapes the crystal.

Using Eq. 3 and considering that, for definition, $LY \equiv L/E = \beta \epsilon k$, we obtain:

$$LY = \tan(\theta_{fit}) \alpha k \quad (5)$$

Now we have to note that the energy calibration we adopted for the Heat axis (as discussed in Sec. 2) implies $H \equiv E$. This means (see Eq. 2) that $\alpha = 1/(1 - k_\gamma)$, being k_γ the absolute LY for γ/β events. Using this last relation in the last equation we finally get for γ/β

$$k_\gamma = \frac{LY_\gamma}{\tan(\theta_{fit}) + LY_\gamma} \quad (6)$$

Using the values of Tab. 1 we get $k_\gamma = 4.6\%$

As a final remark, we point out that the large spread in the light channel (6.8% at the 2615 keV scintillation signal), as observed in the scatter plots, cannot be dominated by the fluctuation of the Poissonian statistics of the absorbed photons. In fact we have that the mean energy absorbed by the LD for a 2615 keV interaction in $CdWO_4$ is of the order of $2.615 \cdot 17.6 = 46$ keV. Assuming ≈ 3 eV/photon we get (assuming a Fano factor=1) the variation in the light channel to be of the order of $1/\sqrt{46000/3} = 0.8\%$. Moreover in other tests performed with different $CdWO_4$ crystal samples (showing “only” a slightly larger LY) we observed the same anticorrelation (in terms of θ_{fit}) but with a *much smaller* spread. Thus the magnitude of this spread is probably dominated by fluctuation in the *overall* light collection efficiency, mainly due to variation of the LY due to position dependent effects. In particular we believe that the grade of the surfaces of the crystal (in terms of diffusion of the scintillation light) plays an important role in this mechanism.

6. Heat absolute scale and Heat Quenching Factor

From the previous sections is evident that for scintillating bolometers the Heat/Energy scale is different for different kinds of interacting particles (or, better, for particles characterized by different LY's). In other words, with our convention for the Heat axis calibration, H measures the total energy of a γ/β particle while it doesn't for an α : the energy calibration of α -particles has to be dealt in a separate way.

This “displacement” for the α lines is evident in our experimental data. In particular the ^{180}W - α line (as can be seen in Fig. 2), whose nominal Q-value is 2516 keV [24], appears (in the γ -calibrated spectra) at 2627 ± 2 keV. This line is rather peculiar since it definitely proves that the mis-calibration cannot

be ascribed to a wrong extrapolation of the detector amplitude vs. energy calibration (the highest γ -line used for the calibration of the Heat signal is indeed the ^{208}Tl line at 2615 keV).

The observed shift in keV, $\Delta E = E_\alpha - E_{\gamma/\beta}$, between the experimentally reconstructed energy and its nominal value is given - for the three mentioned α 's - by:

$$\Delta E_{180W} = 110 \pm 3 \quad \Delta E_{238U} = 172 \pm 3.5 \quad \Delta E_{234U} = 180 \pm 2 \quad (7)$$

This shift can be easily recognized by looking at Fig. 3. Let us consider the case in which the crystal does not scintillate: in this case the two particles, $\alpha 1$ and $\gamma 1$ will show, as example, $E(\alpha 1) < E(\gamma 1)$. If the crystal scintillates and the light escapes the crystal, then the two events will move left on the two different iso-energy lines up to reach the case in which $H(\alpha 1) = H(\gamma 1)$. This means that the energy shift can be written as $\Delta E = H_{\gamma 1}^{L=0} - H_{\alpha 1}^{L=0}$. By simple geometric consideration we get

$$E_{\gamma/\beta}^{Light=0} = E + \frac{Light_{\gamma/\beta}}{\tan(\theta_{fit})} = E \left(1 + \frac{LY_{\gamma/\beta}}{\tan(\theta_{fit})} \right) \quad (8)$$

So, finally, we obtain

$$\Delta E = E [MeV] \cdot \left(\frac{LY_{\gamma/\beta} - LY_\alpha}{\tan(\theta_{fit})} \right) [keV] \quad (9)$$

This formula represents the energy displacement that characterizes two isoenergetic particles impinging a generic scintillating crystal, induced by their different scintillation yield (in our case $LY_{\gamma/\beta}$ and LY_α).

Inserting the experimental values we get:

$$\Delta E_{180W} = 107 \pm 3 \quad \Delta E_{238U} = 175 \pm 6 \quad \Delta E_{234U} = 195 \pm 7 \quad (10)$$

These last values are in very good agreement with the experimental ones in Eq. 7, confirming our model.

To conclude this section, we need to face an important question about the definition of the Heat Quenching Factor (HQF). The conventional way to treat differences in the response of a detector with respect to different type of interacting particles is to introduce relative quenching factors. This is exactly what is done for the scintillation light. In the same way, following what is often done in literature [25, 26, 27, 28], it is possible to introduce the HQF as the ratio between the Heat signal of an α particle and that of a γ/β of the same energy, fully absorbed in the detector. With this definition, for our detector we can quote the HQF measured for the ^{180}W α decay as: $HQF_{180W} = 2627/2516 = 1.044$. It is straightforward to show that in our framework we can easily evaluate $HQF = 1 + \frac{LY_{\gamma/\beta} - LY_\alpha}{\tan(\theta_{fit})} = 1 + \frac{LY_{\gamma/\beta}}{\tan(\theta_{fit})} (1 - QF_\alpha)$.

Even if the definition of the (scintillation) QF and the HQF is identical, due to the different physical mechanism that generates them, they show a *substantial* difference. While the former depends, substantially, only on the “nature” of the scintillator, the HQF, instead, depends on the amount of light (energy) that escapes the crystal: it depends on the size and/or on the quality of the crystal. It is also obvious that since some (minor) amount

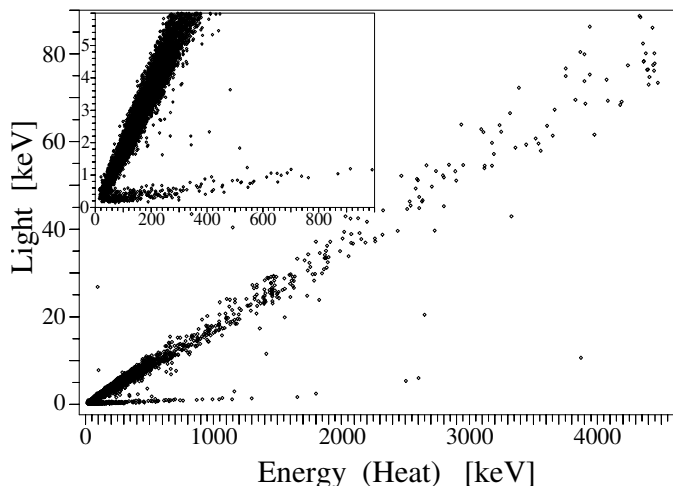


Figure 5: Scatter plot Heat vs. Light obtained in a 8 hour measurement with an Am-Be neutron source. In the inset the low energy region. The energy spectrum extends up to 9 MeV. The 4.44 MeV “line” due to the $^{12}\text{C}^*$ is visible. The intrinsic energy spread of this peak is rather large ($\approx 1\%$) due to Doppler shift induced by the “in flight” decay of $^{12}\text{C}^*$ in the source.

of light that escapes the crystal can be reflected by the reflecting foil and can be absorbed by the crystal itself, then the HQF will depend also by the setup itself.

With the $3 \times 3 \times 2 \text{ cm}^3$ CdWO_4 crystal measured in [9], we found the position of the ^{180}W line at 2670 keV, giving therefore $\text{HQF}=1.061$.

Now one relevant question that could be addressed is: do we have blind channels in our detector? We don’t have a general method to investigate this point, but the results presented so far are still valid even if we assume that a fixed fraction of the total energy is lost in blind channels (we used only Heat/Light ratios in our framework). Then we can investigate if this fraction is different between a γ/β and an α . In other words we could ask ourselves if the sum of the heat and light signals is the same for particle converting the same total energy in the scintillator. We provide therefore an alternative definition of HQF as the ratio of $E_{\alpha}^{\text{Light}=0}$ and $E_{\gamma/\beta}^{\text{Light}=0}$.

Provided that our model that compensates the energy losses due to scintillation is correct, using the values of Eq. 10 we can evaluate the (corrected) HQF:

$$\text{HQF}(^{180}\text{W})=1.001 \pm .002 \quad \text{HQF}(^{238}\text{U})=1.000 \pm .002 \quad \text{and} \\ \text{HQF}(^{234}\text{U})=0.997 \pm .002.$$

These HQF’s are compatible with 1 meaning that there are no intrinsic differences in the heat signal generated by an α with respect to a signal generated by a γ/β ¹. Alternatively we can state that any eventually present blind channel have the same behaviour for the two.

¹We note here that although we always discuss of α ’s it is true that we are always considering the system (α + nuclear recoil), where however the nuclear recoil transports only a small fraction of the total energy.

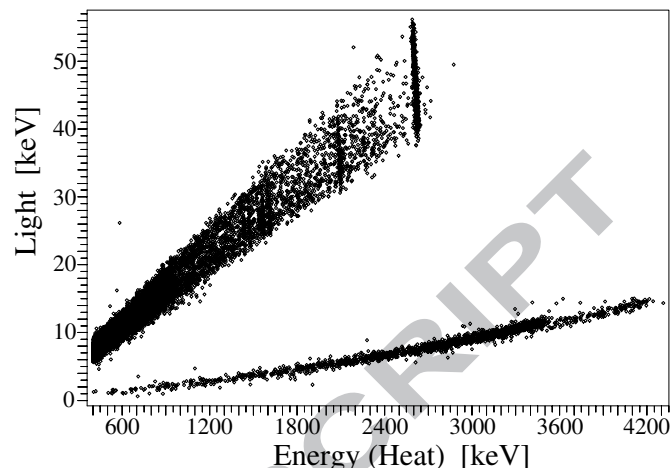


Figure 6: Scatter plot Light vs. Heat obtained in a 93 hours ^{232}Th calibration measurement with a smeared α source facing the CdWO_4 . The liquid standard also contains a small amount of ^{234}U . The alpha curve is completely separated from the γ/β line. Moreover it can be noted that the α events are not belonging to a straight line, showing a decrease of the QF with the energy, as discussed in Sec. 5.

7. Background Rejection: α ’s and neutrons

The detector was also exposed to a neutron source. It consists of a 185 kBq Am-Be source with a neutron production rate of ≈ 10 n/s. The neutron spectrum has its maximum at ~ 5 MeV with an high energy tail reaching 10 MeV. The source was placed inside the shielding of the cryostat, in the same place where the γ sources are inserted. The scatter plot is shown in Fig. 5: neutron direct interactions are clearly visible, especially in the inset. The γ/β events extends well above 2615 keV due to (n,γ) reaction on the surrounding materials but, mainly, by the source itself: for each neutron produced there is 60 % probability to produce an excited state of ^{12}C that emits a γ of 4.44 MeV.

Moreover it has to be remarked (see also Sec. 8) that ^{113}Cd has a huge neutron capture cross section, with a Q-value larger than 9 MeV, so that “mixed events” are possible (a neutron scatters on the Oxygen of the crystal and then is absorbed by the ^{113}Cd with subsequent de-excitation of the nucleus). From Fig. 5 we evaluate the neutron scintillation QF with respect to γ/β to be (0.14 ± 0.03) . The error is dominated by the systematics induced by the extreme weakness of the scintillation light in the region $0 \div 300$ keV, where most of the events are recorded.

In a subsequent test, an α source was mounted close to the detector in order to evaluate its rejection capability. The custom source was build up using $2 \mu\text{L}$ of a 0.1 % ^{238}U liquid Standard. The U “droplets” were dried on a 1 cm^2 Al tape and then covered with a $6 \mu\text{m}$ Aluminized Mylar foil in order to “smear” the alpha energy down to the region of interest. The source was then faced to the crystal on the face opposed to the LD. In Fig. 6 we present the scatter plot obtained with the above mentioned source during a 93 hours ^{232}Th calibration measurement. Fixing at 4σ the acceptance on the light signal of the 2615 keV γ -line, we evaluate a rejection factor for α particles $> 35\sigma$ at the

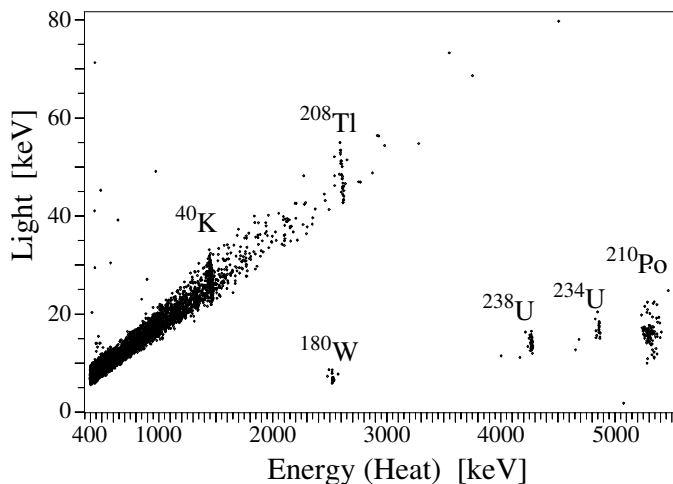


Figure 7: Scatter plot Light vs. Heat obtained in a 433 hours live time background measurement. Part of the ^{210}Po “bump” is characterized by smaller light emission. This could arise from a worse light collection efficiency, especially from the lateral part of the crystal. The energy scale of the α is corrected according with what exposed in Sec. 6.

same energy. As far as the fast neutron interaction is concerned, the rejection factor will be obviously larger, but the most dangerous contribution will arise from thermal neutron absorption through ^{113}Cd . But, at least in principle, thermal neutron can be effectively shielded while fast neutron, via μ -spallation within the shielding or close to it, cannot.

8. Background measurement: internal contaminations

In Fig. 7 we present the Light vs. Heat scatter plot obtained in the 433 h live time background measurement. It can be easily recognized that we observe only two transitions belonging to the Uranium chain, namely ^{238}U and ^{234}U , corresponding to an internal contamination of $3.1 \cdot 10^{-12}$ g/g and $5.7 \cdot 10^{-17}$ g/g, respectively. The Uranium chain is broken at ^{234}U , while a contamination of ^{210}Po is clearly seen. As it often happens, it is rather difficult to determine if this Po contamination arise from its parent, ^{210}Pb . Moreover the absence of a clear peak indicates that the contamination is partially penetrating (few μm) the surface of the crystal (or its surroundings). No events can be ascribed to the Thorium chain, giving a limit of $9 \cdot 10^{-13}$ g/g (95% CL) for ^{232}Th , one of the most “dangerous” contaminant for DBD searches. But the main (unexpected) feature resulting from Fig. 7 is the γ/β background above 2615 keV. These 11 events cannot arise from external γ 's due to ^{208}Tl or ^{214}Bi since the observed background lines (583 keV, 2615 keV and 1764 keV) are too weak to allow such contribution at high energy. The same holds for internal contaminations belonging to Th and U chain. Unrecognized pile-up of the ^{113}Cd - β decay ($Q=320$ keV) with a 2615 keV γ cannot contribute at this level (we expect ~ 2 ev/year). Internal contaminations due to rare high-energy β emitters like ^{106}Rh ($Q=3541$ keV) are extremely difficult to evaluate. But, in any case, we observe γ/β events up to 4.5 MeV, extremely large to be associated with

standard “known” β emitters. On the other hand we previously tested different CdWO_4 scintillating crystals [30], obtaining, with larger statistics, no events above 2615 keV. The main difference with respect to [30] is that in the present work the environmental neutron shielding (consisting of 7 cm of polyethylene and 1 cm of CB_4) was removed from the top of the cryostat, leaving ~ 0.6 m² opening, corresponding to $\approx 8\%$ of the total neutron shielding. We think that the thermal neutron absorption by ^{113}Cd is the most probable mechanism in order to explain the anomalous background above 2615 keV.

9. Conclusions

For the first time a “large” scintillating bolometer to be used for future DBD searches was fully characterized in details in terms of energy resolution, internal contaminations and particle identification capabilities. It was shown how the use of the anticorrelation between light and heat improves the energy resolution by a factor 1.26 @ 88 keV and up to a factor 2.6 @ 2615 keV. We developed, for the first time, a procedure that evaluates the Heat Quenching Factor corrected for the loss of the scintillation light.

This CdWO_4 , grown without any precaution in terms of radiopurity, shows extremely low trace contaminations in U and Th. Nonetheless a future experiment based on this compound should use isotopic enrichment, for two reasons: i) increase the mass of ^{116}Cd ; ii) decrease the fraction of ^{113}Cd in order to avoid β decay and, more important, decrease (n, γ) reaction due to ^{113}Cd .

As a final remark we point out that such a crystal, depleted in ^{113}Cd , could be an extremely interesting compound not only for DBD experiments but also for Dark Matter searches, due to its “uncommon” radiopurity. Our LD, obviously, does not permit to discriminate charged particles with respect to neutral particles below 40 keV. On the other hand, the LY we measured is comparable with that one of the scintillators (CaWO_4 , ZnWO_4) used in the CRESST [31] Experiment, opening, at least in principle, the possibility of having a cryogenic experiment both for DBD and DM Searches.

10. Acknowledgements

The results reported here have been obtained in the framework of the Bolux R&D Experiment funded by INFN, aiming at the optimization of a cryogenic DBD Experiment for a next generation experiment. Thanks are due to E. Tatananni, A. Rotilio, A. Corsi and B. Romualdi for continuous and constructive help in the overall setup construction. Finally, we are especially grateful to Maurizio Perego for his invaluable help in the development and improvement of the Data Acquisition software.

References

- [1] V. I. Tretyak and Yu. G. Zdesenko, Atomic Data and Nuclear Data Tables, **80** (2002): 83;
- [2] S. Elliott and P. Vogel, Ann. Rev. Nucl. Part. Sci. **52** (2002):115;

- [3] A. Morales and J. Morales, Nucl. Phys. B (Proc. Suppl.) **114** (2003):141;
- [4] F. T. Avignone III, S. R. Elliott and J. Engel (2006), Rev. Mod. Phys., **80** (2008):481.
- [5] S. Pirro, Eur. Phys. J. A **27**, S1 (2006):25
- [6] C. Arnaboldi, *et al.*, Phys. Rev. C **78** (2008):035502
- [7] C. Arnaboldi, *et al.*, Nucl. Instr. and Meth. A **518** (2004):775
- [8] M. Pavan, *et al.*, Eur. Phys. J A **36** (2008):159
- [9] S. Pirro, *et al.*, Physics of Atomic Nuclei **69** No.12 (2006):2109
- [10] F.A. Danivich, *et al.*, Phys. Rev. C **68** (2003):035501
- [11] G. Bellini, *et al.*, Phys. Lett. B **493** (2000):216
- [12] J. Argyriades *et al.*, Nucl. Instr. and Meth. A **606** (2009):449
- [13] J. McGrath *et al.*, Nucl. Instr. and Meth. A **615** (2010):57
- [14] S. Pirro, *et al.*, Nucl. Instr. and Meth. A **559** (2006):361
- [15] A. Alessandrello, *et al.*, Nucl. Instr. and Meth. A **412** (1998):454
- [16] C. Arnaboldi, G. Pessina, E. Previtali, IEEE Tran. on Nucl. Sci. **50** (2003):979
- [17] S. Pirro, Nucl. Instr. and Meth. A **559** (2006):672
- [18] C. Arnaboldi, G. Pessina, S. Pirro, Nucl. Instr. and Meth. A **559** (2006):826
- [19] C. Arnaboldi, *et al.*, Nucl. Instr. and Meth. A **520** (2004):578
- [20] E. Conti, *et al.*, Phys Rev. B **68**(2003):054201
- [21] J. Amare, *et al.*, Appl. Phys. Lett. **87** (2005):264102
- [22] V.I. Tretyak, Astropart. Phys. **33** (2010):40
- [23] J.B. Birks, Proc. Phys. Soc. A **64** (1951):874
- [24] C. Cozzini, *et al.*, Phys Rev. C **70**(2004):064606
- [25] A. Alessandrello, *et al.*, Phys Lett B **408**(1997):465
- [26] N. Coron, *et al.*, Phys Lett B **659**(2008):113
- [27] G. Angloher, *et al.*, Astropart. Phys. **23** (2005):325
- [28] A. Benoit, *et al.*, Nucl. Instr. and Meth. A **577** (2007):558
- [29] N. Coron, *et al.*, J. Low Temp. Phys. **151** (2008):865
- [30] L. Gironi, C. Arnaboldi, S. Capelli, O. Cremonesi, G. Pessina, S. Pirro and M. Pavan, Opt. Mat. **31** Issue 10 (2009):1388
- [31] G. Angloher *et al.*, Astropart. Phys., **31** (2009):270