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Olivier Chesneau, A. Meilland, D. P. K. Banerjee, Jean-Baptiste Le Bouquin, H. Mcalister, et al.. The 2011 outburst of the recurrent nova T Pyx. Evidence for a face-on bipolar ejection.. Astronomy and Astrophysics - A&A, 2011, pp.__. hal-00625395v2

HAL Id: hal-00625395

<https://hal.science/hal-00625395v2>

Submitted on 22 Sep 2011 (v2), last revised 26 Sep 2011 (v3)

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LETTER TO THE EDITOR

The 2011 outburst of the recurrent nova T Pyx. Evidence for a face-on bipolar ejection.

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Received, accepted.

ABSTRACT

Aims. T Pyx is the first recurrent nova historically studied, seen in outburst six times between 1890 and 1966 and then not for 45 years. We report on near-IR interferometric observations of the recent outburst of 2011. We compare expansion of the H and K band continua and the Bry emission line, and infer information on the kinematics and morphology of the early ejecta.

Methods. We obtained near-IR observations of T Pyx at dates ranging from $t=2.37d$ to $t=48.2d$ after the outburst, with the CLASSIC recombiner, located at the CHARA array, and with the PIONIER and AMBER recombiners, located at the VLTI array. These data are supplemented with near-IR photometry and spectra obtained at Mount Abu, India.

Results. Slow expansion velocities were measured ($\leq 300 \text{ km s}^{-1}$) before $t=20d$ (assuming $D=3.5 \text{ kpc}$). From $t=28d$ on, the AMBER and PIONIER continuum visibilities (K and H band, respectively) are best simulated with a two component model consisting of an unresolved source plus an extended source whose expansion velocity onto the sky plane is lower than $\sim 700 \text{ km s}^{-1}$. The expansion of the Bry line forming region, as inferred at $t=28d$ and $t=35d$ is slightly larger, implying velocities in the range $500\text{-}800 \text{ km s}^{-1}$, still strikingly lower than the velocities of $1300\text{-}1600 \text{ km s}^{-1}$ inferred from the Doppler width of the line. Moreover, a remarkable pattern was observed in the Bry differential phases. A semi-quantitative model using a bipolar flow with a contrast of 2 between the pole and equator velocities, an inclination of $i=15^\circ$ and a position angle $\text{P.A.}=110^\circ$ provides a good match to the AMBER observables (spectra, differential visibilities and phases). At $t=48d$, a PIONIER dataset confirms the two component nature of the H band emission, consisting of an unresolved stellar source and an extended region whose appearance is circular and symmetric within error bars.

Conclusions. These observations are most simply interpreted within the frame of a bipolar model, oriented nearly face-on. This finding has profound implications for the interpretation of past, current and future observations of the expanding nebula.

Key words. Techniques: high angular resolution; (Stars:) novae, cataclysmic variables; individual: T Pyx; Stars: circumstellar matter

1. Introduction

A classical nova eruption results from a thermonuclear runaway on the surface of a white dwarf which is accreting material from a companion star in a close binary system. T Pyxidis (T Pyx) is a unique recurrent nova that was in outburst six times between 1890 and 1966 (intervals of $\sim 20 \text{ yr}$). T Pyx was discovered in outburst at a visual magnitude of 13.0 on 2011 April 14.29 UT (JD=2455665.79); which we take as $t_0=0$ (Waagan et al. 2011). This is the first outburst of T Pyx since December 7, 1966, nearly 45 years before.

The evolution of the nova is relatively slow, thereby providing time and scope for organizing joint observations with optical

interferometry arrays such as CHARA and the VLTI. T Pyx is surrounded by an interesting nebula in expansion that has been investigated by the HST during more than 10yrs (Schaefer et al. 2010, and references therein). The knots are expanding in the plane of the sky with velocities ranging from roughly 500 to 715 km s^{-1} . In contrast, the velocities inferred from Doppler widths of the ejecta of recent outbursts were observed to be much faster at about 1500 km s^{-1} . Although T Pyx is a well-observed system, it still has many mysteries. Why did the ejecta expand so slowly in the plane of the sky? An important spectroscopic study of the binary system from Uthas et al. (2010) provided evidence of a low-inclination for the system orbit ($i=10\pm 2^\circ$), a particularly important constraint for the interpretation of interferometric data, as it appears that the ejecta emitted around these outbursting sources are rarely spherical.

This letter presents optical interferometry measurements obtained from different facilities which provide important infor-

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* Based on observations made with CHARA at Mount Wilson observatory and the VLTI at Paranal Observatory under program 287.D-5012, 287.D-5023, 087.C-0702

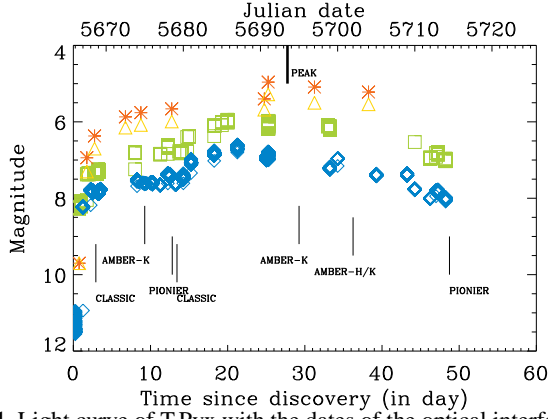


Fig. 1. Light curve of T Pyx with the dates of the optical interferometry observations. Blue diamonds indicate a subset of AAVSO data in V, and green squares in I. Orange triangles and red stars indicate H and K band photometry from Mt Abu (India) .

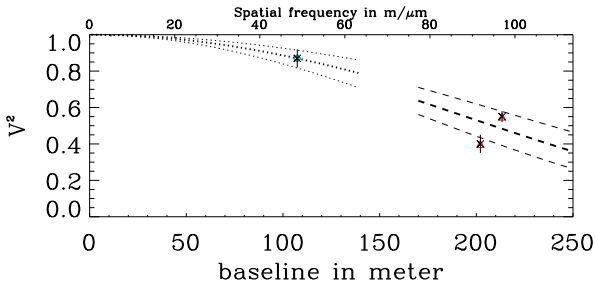


Fig. 2. K-band interferometric visibilities obtained with CLASSIC at $t=2.92$ (green diamond) and $t=13.93$ (red triangles). The thick dotted and dashed lines indicate are the UD curves corresponding to Table 2.

mation when included within a common frame of interpretation. The observations are presented in Sect. 2. In Sect. 3 we analyze the continuum measurements by means of simple geometrical model, and the differential observables through the Bry line using a simple model and then discuss the results in Sect. 4.

2. Observations

Near-infrared JHK photometric and spectroscopic observations were obtained on a regular basis from the 1.2m telescope at the Mt. Abu Observatory, India. These measurements helped to prepare the interferometric observations and to evaluate the relative contribution of the various continuum and line components (Fig.1). Initial observations are reported in Banerjee & Ashok (2011) while a fuller study is in preparation.

Prompt *broad-band interferometric observations* were secured with CLASSIC, a two-telescope high sensitivity system located at CHARA on Mt. Wilson (ten Brummelaar et al. 2005). Despite the faintness and low declination of the source, observations in the K-band were obtained at $t=2.92$ d ($K\approx 6.4$, from Mt Abu observations) and $t=13.93$ d ($K\approx 5.7$). The log of the observations is presented in Table 1 and the data in Fig.2.

Several *interferometric observations at medium spectral resolution* ($R=1500$) across the Br- γ line were obtained with AMBER, a 3-telescope combiner located at the VLTI (Petrov et al. 2007). The first observations were performed with the 1.8m Auxiliary Telescopes (ATs) at $t=8.81$ d, when the source was below the interferometric sensitivity limit of AMBER ($K\approx 5.7$), but a useful spectrum was obtained. The second and third measurements, obtained with the 8.2m Unit Telescopes (UTs) at

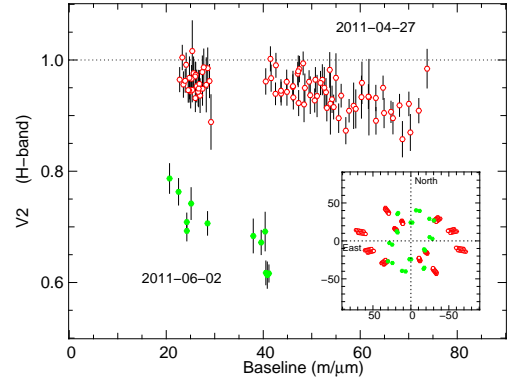


Fig. 3. H-band interferometric visibilities obtained with PIONIER at $t=12.81$ (open red) and $t=48.74$ (filled green). The corresponding uv-plane is displayed in the subpanel.

$t=28.76$ ($K\approx 4.9$) and $t=35.77$ d ($K\approx 5$), provided good quality dispersed visibilities, closure and differential phases (see Fig.5). Unfortunately, the calibrator measurement for the last date is of poor quality preventing any reliable calibration of the absolute visibility.

Imaging broad-band interferometric observations were obtained at $t=12.81$ d ($H\approx 6$) with the PIONIER visitor instrument (Berger et al. 2010; Le Bouquin et al. 2011). These observations provided the simultaneous measurement of 6 absolutely calibrated visibilities and 4 closure-phases in the H-band, therefore allowing the study of the spatial morphology of the near-infrared emission. A critical second observation was obtained at $t=48.74$ d ($H\approx 6$), again with the ATs (Fig.3).

3. Analysis

The absolute visibility measurements were fitted with simple geometrical models using the LITpro software (Tallon-Bosc et al. 2008, JMMC). The results are shown in Table 2. A simple uniform disk (UD) model, i.e. a circular disk of uniform brightness in the plane of the sky, was fitted to the measurements for the early observations. For later observations, a two component model consisting in an unresolved component, and a co-centered uniform disk was used. No evidence of asymmetry was detected in the data.

The first K-band CLASSIC measurement, obtained with a ~ 100 m baseline was consistent with a weakly resolved source, while the second set ($t=13.93$ d) with a baseline about twice as long, shows a resolved source whose size does not seem to be dramatically changed. The first PIONIER measurements ($t=12.81$ d) provided a H-band UD diameter significantly smaller than the CLASSIC K band measurements, even taking into account the small time difference between the two measurements. This effect cannot be attributed to emission lines seen at these dates, that contribute less than 10-15% of the flux in the H band and less than 5% in the K band. A large scale component with a rising flux contribution in the K band may account for the observations, and is consistent with a H-K flux difference of 0.33 mag that is observed consistently during the event. Assuming a two component model, with a fixed contribution of $70\pm 10\%$ from a 'stellar' source with a diameter of 0.6mas, the extended K-band source should have a diameter larger than 1.5-2mas, and therefore be almost fully resolved by CLASSIC. Assuming a distance of $D=3.5\pm 1$ kpc from Schaefer et al. (2010), that may be a lower limit (Shore et al. 2011), the expansion velocity inferred from such an extended component is about 500km s^{-1} , while the

Table 2. Analysis of the V^2 using geometrical models.

Instrument	Spectral Band	t-t ₀ day	Single component model UD diameter (mas)	Double components model		
				Unres. Flux (%)	UD Flux (%)	UD diam. (mas)
CHARA/CLASSIC	broad K	2.92	1.±0.2	-	-	-
VLTI/PIONIER	broad H	12.81	0.6±0.1	-	-	-
CHARA/CLASSIC	broad K	13.93	1.12±0.14	-	-	-
VLTI/AMBER	2.1±0.05μm	28.76	2.58±0.3	65±12%	35±8%	7.3±0.3
VLTI/PIONIER	broad H	48.74	2.23±0.1	83±9%	17±2%	8.5±0.2

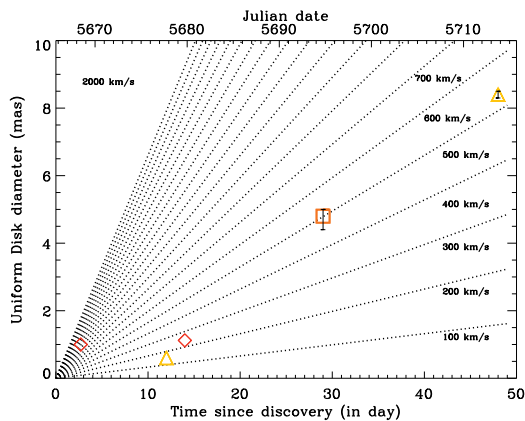


Fig. 4. Result of the Uniform Disk (UD) estimates from the continuum V^2 measurements from the various interferometers. Before $t=20$ d, the source size is estimated using a single UD in the H and K bands (PIONIER: orange triangle; CLASSIC: red diamonds). The last points indicate the extended source in the double component model (AMBER: red square; PIONIER: orange triangle)

H-band core expansion is estimated to be $\sim 100 \text{ km s}^{-1}$ (Fig. 4). Interestingly, the FWHM of the $\text{Br}\gamma$ measured from the AMBER spectrum at $t=8.81$ d is 590 km s^{-1} , i.e. is consistent with our hypothesis of the expanding extended component (Fig. 5).

Then follows a second epoch during which AMBER interferometric data were obtained on two dates, yet providing calibrated visibilities only at $t=28.76$ d. A single UD does not account well for the observed visibilities at $t=28.76$ d (reduced $\chi_r^2=35$), the two component model provides a better fit to the data ($\chi_r^2=5$). This gives an upper limit for the expansion of the extended continuum component of 700 km s^{-1} . This is in contrast with the width of the $\text{Br}\gamma$ line, for which we measured a FWHM of $1050 \pm 50 \text{ km s}^{-1}$. Furthermore, the Doppler velocity associated with the P-Cygni absorption in the line is found to be $-1450 \pm 100 \text{ km s}^{-1}$ (Fig. 5). The $\text{Br}\gamma$ line-forming region should therefore expand much faster, and a large visibility drop should be measured through the $\text{Br}\gamma$ line, as seen for RS Oph (Chesneau et al. 2007). However, the dispersed $\text{Br}\gamma$ visibilities are only slightly lower than the nearby continuum, implying a moderate diameter increase of less than 10%. One week later, the visibilities dropped in the line indicating a large expansion of the $\text{Br}\gamma$ line-forming region (taking into account the 25% increase of the line flux). At that time the line FWHM was measured to be $1600 \pm 50 \text{ km s}^{-1}$ and the P-Cygni absorption indicated a wind velocity of $1800 \pm 100 \text{ km s}^{-1}$. The differential phases show a complex structured signal that can be described by 2 opposite S-shaped signals, with variations related to the baseline lengths and P.A. dependency. The pattern is symmetrical about the line center, with the width and amplitude of the signal increasing between the two dates, although the amplitude never went beyond 10° (again by contrast with what was observed for RS Oph).

The last observations, performed at $t=48.74$ d with PIONIER, bring complementary and crucial information, owing to the larger uv coverage involved. The H band source is now well resolved and departs from a simple model. A good fit ($\chi^2 = 1.1$) is reached using the two component model (Table 2). The extended component estimated expansion velocity in the plane of the sky is lower than 700 km s^{-1} . The closure phases do not exceed 2.5° . Moreover using a flattened structure for the extended component does not improve the fit and constrains the aspect ratio to 1 ± 0.07 . This implies that the complex yet weak phase signal seen by AMBER originates from a source with a predominantly symmetrical appearance.

4. A face-on bipolar event

The heterogeneous data set originating from different instruments provides intriguing data, and is unique in view of the complementary constraints provided for the analysis. The interpretation of these data can be divided into several key temporal steps.

- CLASSIC and PIONIER data obtained at $t=2.92$ d, 12.81d and $t=13.93$ d show an extended H and K band source. The fact that T Pyx is resolved so early is intriguing. An hypothesis might be a light echo witnessing a close-by circumbinary environment, adding incoherent flux to the measurement. The expansion rate is then low between the first and second measurements, and a striking difference is observed between the H and the K band inferred diameters. This suggests an advanced decoupling between a shrinking optically thick core and an expanding free-free emitting optically thin envelope.

- The AMBER data obtained at $t=28.76$ d provide evidence that the $\text{Br}\gamma$ line forming region projected onto the sky is formed close to the expanding continuum, while being more extended at $t=35.77$ d. The Doppler velocity accelerated from 1300 km s^{-1} to 1600 km s^{-1} in the time interval. A striking differential phase pattern is observed.

- The PIONIER dataset secured at $t=48.74$ d is critical for establishing the two component nature of the emission, consisting in an unresolved stellar source and an extended region whose appearance is circular and symmetric within error bars.

A face-on bipolar event could account for the ensemble of information described above, in accordance with the finding of a low inclination for the system (Uthas et al. 2010). In particular, the differential phase pattern can be linked to the geometry and kinematics of the ejecta. We developed a ‘toy’ model that provides a good match to the observations. Given two ad-hoc three-dimensional distributions, one for the ‘emission’ of the ejecta and one for the velocity field, we reconstructed intensity maps in narrow spectral bands in the emission line and then computed the corresponding visibilities and differential phases. The intensity map is created considering that the matter was ejected during a brief outburst (best ascribed by a shell), propagating at $v_r(\theta)$.

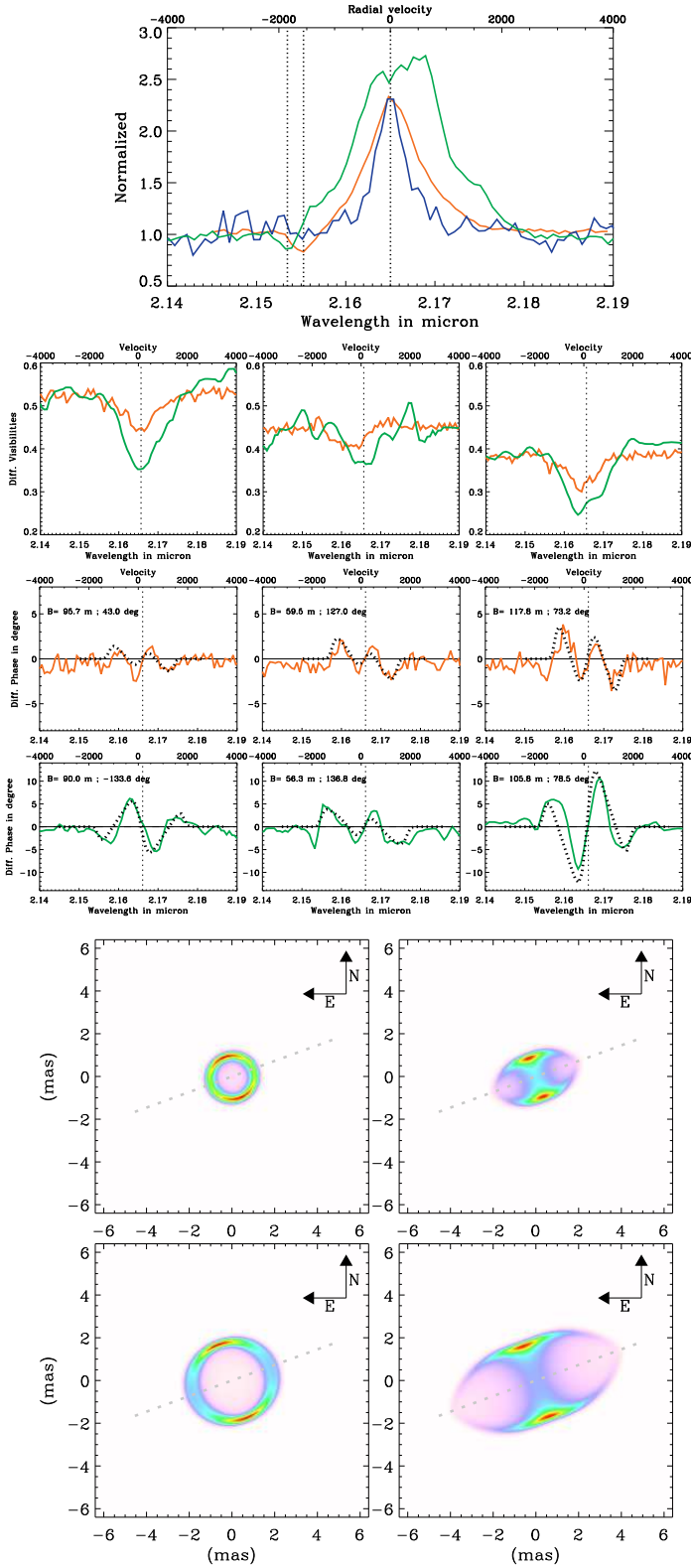


Fig. 5. AMBER data. **Top:** comparison between the Bry line at $t=8.81$ (blue curve), $t=28.76d$ (red curve) and $t=35.77d$ (green curve). **Middle:** Above - differential visibility comparison between $t=28.76d$ and $t=35.77d$ (scaled to the continuum V^2 at $t=28.76d$). Below: same with differential phases. The phases are compared with the phases from the model (dashed line, $\chi^2=1.1$ and 1.4 , respectively). **Bottom:** Bry Bipolar-flow model (without central source) seen at $i=90^\circ$ (right) and $i=10^\circ$ (left, the best model) and P.A.= 110° (best model).

Consequently, the geometry is directly related to kinematics of the ejecta. We used the following radial expansion law:

$$v_r(\theta) = v_{pole} + (v_{eq} - v_{pole}) \sin(\theta) \quad (1)$$

where θ is the colatitude, and v_{pole} and v_{eq} are the polar and equatorial radial velocities, respectively.

We also considered an emission decreasing according to a power law of the distance. At a given epoch t the 3D intensity distribution is proportional to:

$$I(r, \theta, \phi) \propto \frac{1}{r^\alpha} \exp\left[-\frac{(v_r t - r)^2}{2\sigma_r^2}\right] \quad (2)$$

Using this model we were able to fit the Bry differential visibilities and phases, as well as the line profile for the two epochs. The parameters of the best model are: $i=15^\circ$, P.A. of the polar axis of 110° , $\alpha=2$ and the velocities for the two epochs:

$$-v_{pole}=1200 \text{ km s}^{-1} \text{ and } v_{eq}=600 \text{ km s}^{-1} \text{ at } t=28.76d$$

$$-v_{pole}=1600 \text{ km s}^{-1} \text{ and } v_{eq}=700 \text{ km s}^{-1} \text{ at } t=35.77d$$

The fit of the differential phases at the two epochs and the model images are presented in Fig 5. The polar and equatorial velocities are in good agreement with the Doppler and sky plane velocities estimated in Sect.3. Furthermore, the P.A. of the equatorial plane overdensity is oriented in a direction similar to the P.A. of the faint X-ray nebula (Balman 2010).

The face-on bipolar nebula allows one to better understand the curious nebula scrutinized with HST (Schaefer et al. 2010; Shara et al. 1997, 1989). The knots are concentrated in a ring ($3.2-6''$), expanding radially with a velocity in the restricted range of $500-700 \text{ km s}^{-1}$, and with a mean radial velocity of about 500 km s^{-1} (O'Brien & Cohen 1998). Deciphering between a projected sphere and a bipolar structure producing a dense, face-on ring is difficult, considering that radial velocity measurements of individual clumps are missing.

Some recent examples suggest that bipolarity in the ejecta of classical/recurrent novae may be relatively frequent: RS Oph (Ribeiro et al. 2009; Bode et al. 2007; Chesneau et al. 2007), V445 Pup (Woudt et al. 2009) or V1280 Sco (Chesneau et al. 2008, Chesneau et al., in prep.). A significant difference though exists between the T Pyx and RS Oph environments: the lack of material around T Pyx, witnessed for instance by the lack of hard X-rays (Kuulkers et al. 2011), leads us to favor a bipolarity induced by a process internal to the system, whether by the common envelope interaction with the companion, since the development of the event is relatively slow, or by invoking an intrinsically bipolar ejection related to a spun-up central star.

Acknowledgements. The CHARA Array is funded by the National Science Foundation through NSF grant AST-0908253, by Georgia State University, the W. M. Keck Foundation, the Packard Foundation, and the NASA Exoplanet Science Institute. Research at the Physical Research Laboratory is funded by the Dept. of Space, Govt. of India. STR acknowledges partial support from NASA grant NNH09AK731.

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Table 1. Journal of interferometric observations.

date	MJD	t-t ₀	Instrument	Base	projected baselines		Calibrators ^a
	2450000.5+	Phase			length [m]	P.A. [deg]	
2011/04/17	5668.16	2.92	CLASSIC	W1-W2	107.4	-87	HD 78752, HD 79290
2011/04/23	5674.04	8.81	AMBER	K0-A1-I1	45.4/100.3/120.2	-159.0/96.2/ -105.0	HD 73947
2011/04/26	5678.06	12.81	PIONIER	A1-G1-I1-K0	74/94/42/113/82/44	-52/-74/-127/81/39/23	HD 78739
2011/04/28	5679.18	13.93	CLASSIC	W1-E2	202.1/213.3	-79.4/82.8	HD 78752, HD 79290
2011/05/12	5694.00	28.76	AMBER	UT1-3-4	59.5/95.6/117.7	127.1/43.1/73.3	HD 73947
2011/05/20	5701.02	35.77	AMBER	UT1-3-4	56.2/89.9/105.6	136.9/-133/78.6	HD 73947, HD 87303
2011/06/01	5713.99	48.74	PIONIER	D0-G1-H0-I1	68/47/63/67/37/40	171/96/132/25/57/180	HD 78739

^a Calibrator angular diameters from SearchCal@JMMC(Bonneau et al. 2006): HD 78752 (G0V, 0.22±0.02mas), HD 79290 (A0V, 0.13±0.01mas), HD 73947 (K2III, 0.86±0.02mas), HD 87303 (K2III, 0.90±0.07mas), HD 78739 (K0III, 0.32±0.02mas).