



**HAL**  
open science

# Can a primordial black hole or wormhole grow as fast as the universe?

B J Carr, Tomohiro Harada, Hideki Maeda

► **To cite this version:**

B J Carr, Tomohiro Harada, Hideki Maeda. Can a primordial black hole or wormhole grow as fast as the universe?. *Classical and Quantum Gravity*, 2010, 27 (18), pp.183101. 10.1088/0264-9381/27/18/183101 . hal-00625158

**HAL Id: hal-00625158**

**<https://hal.science/hal-00625158>**

Submitted on 21 Sep 2011

**HAL** is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

# Can a primordial black hole or wormhole grow as fast as the universe?

<sup>1,2</sup>B. J. Carr,<sup>\*</sup> <sup>3</sup>Tomohiro Harada,<sup>†</sup> and <sup>4</sup>Hideki Maeda<sup>‡</sup>

<sup>1</sup>*Astronomy Unit, Queen Mary University of London, Mile End Road, London E1 4NS, UK*

<sup>2</sup>*Research Center for the Early Universe, Graduate School of Science, University of Tokyo, Tokyo 113-0033, Japan*

<sup>3</sup>*Department of Physics, Rikkyo University, Tokyo 171-8501, Japan*

<sup>4</sup>*Centro de Estudios Científicos (CECS), Casilla 1469, Valdivia, Chile*

(Dated: June 3, 2010)

This review addresses the issue of whether there are physically realistic self-similar solutions in which a primordial black hole is attached to an exact or asymptotically Friedmann model for an equation of state of the form  $p = (\gamma - 1)\rho c^2$ . In the positive pressure case ( $1 < \gamma < 2$ ), there is no solution in which the black hole is attached to an exact Friedmann background via a sonic point. However, there is a one-parameter family of black hole solutions which are everywhere supersonic and asymptotically *quasi*-Friedmann, in the sense that they contain a solid angle deficit at large distances. Such solutions exist providing the ratio of the black hole size to the cosmological horizon size is above some critical value and they include “universal” black holes with an apparent horizon but no event horizon. Solutions of this kind may also exist in the  $2/3 < \gamma \leq 1$  case. In the stiff case ( $\gamma = 2$ ), there is no self-similar solution in an exact background unless the matter turns into null dust before entering the event horizon; otherwise the only black hole solutions are probably asymptotically quasi-Friedmann universal ones. For a dark-energy-dominated universe ( $0 < \gamma < 2/3$ ), there is a one-parameter family of black hole solutions which are properly asymptotically Friedmann (i.e. with no angle deficit) and the ratio of the black hole size to the cosmological horizon size is *below* some critical value. Above this value, one finds a self-similar cosmological wormhole solution which connects two asymptotic regions, one exactly Friedmann and the other asymptotically quasi-Friedmann. The existence of self-similar black holes in phantom fluids ( $\gamma < 0$ ) or tachyonic fluids ( $\gamma > 2$ ) remains uncertain. We also consider the possibility of self-similar black hole solutions in a universe dominated by a scalar field. This is like the stiff fluid case if the field is massless, but the situation is less clear if the scalar field is rolling down a potential and therefore massive, as in the quintessence scenario. Although no explicit asymptotically Friedmann black hole solutions of this kind are known, they may exist if the black hole is not too large and this implies that a black hole might grow as fast as the cosmological horizon in a quintessence-dominated universe in some circumstances.

PACS numbers: 04.70.Bw, 97.60.Lf, 04.40.Nr, 04.25.Dm, 95.35.+d

## I. HISTORICAL INTRODUCTION

Over the last 40 years there has been much interest in how fast a black hole formed in the early universe would grow. In particular, attention has focussed on whether such a primordial black hole (PBH) could grow as fast as the cosmological horizon, corresponding to a spherically symmetric self-similar solution. The answer turns out to be very dependent on the equation of state of the background universe. This is usually taken to have the form  $p = (\gamma - 1)\rho c^2$ , where  $p$  is the pressure,  $\rho$  is the mass density and  $\gamma$  is a constant, because this is the most general barotropic equation of state compatible with self-similarity. This review is therefore divided into separate sections, each focussing on a different range of the parameter  $\gamma$ . However, we will start with a brief historical account of the problem because this is illuminating.

As first pointed out by Zel’dovich and Novikov [1], a simple Bondi-type accretion analysis in a radiation-dominated Universe ( $\gamma = 4/3$ ) suggests that a PBH would not grow much at all if it were much smaller than the cosmological horizon at formation but that it could

grow at the same rate as the universe if its initial size were comparable to it. (The term “cosmological horizon” should here be interpreted as the Hubble horizon if the PBHs form after an inflationary period but the particle horizon otherwise.) One might expect the latter situation to apply, since a PBH must be bigger than the Jeans length at formation [2], so this suggests that any PBH might grow to the horizon mass at the end of the radiation era, which is around  $10^{17}M_{\odot}$ . Since there is no evidence for such enormous black holes, for a while it was assumed that no PBHs ever formed.

However, the validity of the Zel’dovich-Novikov calculation is questionable when the black hole size is comparable to the horizon size because it neglects the expansion of the universe and is not fully relativistic. Indeed, the conclusion that a PBH could grow at the same rate as the universe in the radiation-dominated era was disproved by Carr and Hawking [3]. They demonstrated this by proving that there is no self-similar solution which contains a black hole attached to an exact flat Friedmann background via a sonic point (i.e. in which the black hole forms by purely causal processes). The Zel’dovich-Novikov prediction is therefore definitely misleading in this case. Since the PBH must soon fall well below the horizon size, when their argument should be valid, this suggests that PBHs would not grow much at all.

This gave the subject of PBHs a new lease of life and motivated Hawking to consider the quantum effects asso-

---

\*Electronic address: B.J.Carr-at-qmul.ac.uk

†Electronic address: harada-at-rikkyo.ac.jp

‡Electronic address: hideki-at-cecs.cl

ciated with black holes (since only PBHs could be small enough for these to be significant). Ultimately, this led to his discovery of black hole radiation [4], so it is ironic that a consideration of PBH accretion led to the conclusion that they evaporate!

Carr and Hawking also claimed that there are self-similar solutions which are *asymptotically* – rather than *exactly* – Friedmann at large distances from the black hole. However, these correspond to special acausal initial conditions, in which matter is effectively thrown into the black hole at every distance; they do not contain a sonic point because they are supersonic everywhere. Indeed, such solutions exist even when the cosmological fluid is pressureless ( $\gamma = 1$ ) [3]. They also found solutions in which the whole universe is inside the black hole; these are now termed “universal” black holes.

Subsequently, the Carr-Hawking analysis was extended to perfect fluids with more general equation of state, though still of the form  $p = (\gamma - 1)\rho c^2$  required by self-similarity. It was proved that the non-existence of self-similar black holes in an exact Friedmann background applies for all values of  $\gamma$  in the range 1 to 2 [5, 6]. Indeed, the only physical self-similar solution which can be attached to an exact external Friedmann solution via a sonic point is Friedmann itself. There are still acausal black hole solutions but these are again supersonic everywhere; the asymptotically Friedmann solutions which reach a sonic point represent density perturbations which grow at the same rate as the universe rather than black holes [7]. Later it was realized that none of these positive-pressure self-similar solutions are strictly asymptotically Friedmann [8]: there is a solid angle deficit at large distances which might in principle show up in the angular diameter test. It would therefore be more accurate to describe them as asymptotically “quasi-Friedmann”.

The attempt to extend the analysis to stiff fluids ( $\gamma = 2$ ) led to some controversy. Lin *et al.* [9] at first claimed that there *is* a self-similar black hole solution in an exact Friedmann background in this case. However, Bicknell and Henriksen [10] showed that this conclusion is invalid because Lin *et al.* had misidentified the point corresponding to the black hole event horizon. Bicknell and Henriksen did construct a numerical self-similar solution containing a black hole but it required the stiff fluid to turn into null dust at some point, which is rather contrived. The only possibility might be universal black hole solutions which are asymptotically quasi-Friedmann.

Later, attention focussed on the growth of a black hole in a universe is dominated by a scalar field, as expected in many cosmological contexts. If the scalar field is massless (i.e. if there is no scalar potential), then one might expect the same result to apply as in the stiff fluid analysis, since it is well known that a scalar field is equivalent to a stiff fluid provided the gradient is everywhere timelike [11]. Indeed, the conclusion that there is no self-similar non-universal black hole solution in an exact or asymptotically Friedmann background dominated by a massless scalar field was supported by both numerical studies [12] and analytical calculations [13].

However, the situation is more complicated if there is

a scalar potential (i.e. if the scalar field is massive), and the discovery that the universe is currently accelerating suggests that this may be the case [14]. This has led to a study of black hole accretion in quintessence-dominated universes. Indeed, an argument similar to that advocated by Zel’dovich and Novikov has resurfaced in this context in order to explain the origin of the  $10^6$  to  $10^9 M_\odot$  black holes thought to reside in galactic nuclei [15]. While there are several scenarios for the formation of such supermassive black holes, one possibility – first proposed by Bean and Magueijo [16] – is that they originated in the early universe and grew self-similarly through accretion of quintessence before cosmological nucleosynthesis, as well as by astrophysical processes later.

Since this proposal is motivated by a Bondi-type argument which neglects the cosmological expansion, it is just as questionable as the original Zel’dovich-Novikov one. This has led to the search for self-similar black hole solutions in quintessence-dominated universes. In this case, the similarity assumption requires that the potential have an exponential form. Our analysis in Ref. [13] then shows that there is no self-similar black hole solution in a wide variety of circumstances. However, it does not exclude such a solution in an accelerating asymptotically Friedmann or quasi-Friedmann background, and this is perhaps the most interesting case.

The acceleration of the universe can also be explained if its density is dominated by a perfect fluid with  $0 < \gamma < 2/3$ , so this has motivated us to look for self-similar black hole solutions in this case [17, 18]. We describe such a fluid as “dark energy” and there is an important difference from the positive-pressure case in that there are now no sound-waves (the sound-speed being imaginary). Quintessence might be regarded as a variant in which the parameter  $\gamma$ , rather than being constant, varies as the scalar field rolls down its potential. However, there are still sound-waves in this case, so we prefer to distinguish between these two scenarios.

The lack of sound-waves in the dark energy case means that one can immediately exclude black hole solutions in an exact Friedmann background. However, as discussed in Refs. [17, 18], there are again asymptotically Friedmann solutions containing black holes. Indeed, unlike the positive-pressure case, these are genuinely asymptotically Friedmann rather than asymptotically quasi-Friedmann and the associated inhomogeneities may arise naturally in the inflationary scenario. There is a one-parameter family of such solutions but (in contrast to the implication of the Zel’dovich-Novikov argument) they only exist if the black hole is not too large compared to the particle horizon. These solutions are not analogous to the positive-pressure solutions which Carr and Hawking originally sought but they might nevertheless be physical.

If one tries to find an asymptotically Friedmann self-similar solution with a black hole which is larger than the upper limit in the  $0 < \gamma < 2/3$  situation, one obtains a cosmological wormhole instead. The transition occurs as the black hole apparent horizon approaches the cosmological apparent horizon, after which both horizons disappear. This is in contrast to the  $2/3 < \gamma < 2$  case,

where the two apparent horizons never merge [5] and one tends to a separate closed universe as the black hole size increases. In the wormhole solution the metric tends to an asymptotically Kantowski-Sachs form as one approaches the wormhole throat, this corresponding to a minimum physical radius, and the solution then connects to another asymptotically Friedmann or asymptotically quasi-Friedmann universe.

This paper provides a comprehensive discussion of all these solutions. For the most part, we will avoid mathematical technicalities, so the number of equations is minimized. Although the paper is intended as a general review of previous work in this area, we believe that bringing all the cases together is illuminating and leads to some original insights. In successive sections, we consider positive-pressure fluids, stiff fluids, scalar and quintessence fields, dark-energy fluids and finally more exotic possibilities (phantom fluids with  $\gamma < 0$ , negative pressure fluids with  $2/3 < \gamma < 1$  and tachyonic fluids with  $\gamma > 2$ ). We conclude that there are certainly self-similar asymptotically Friedmann black hole solutions in the  $0 < \gamma < 2/3$  case and there may also be in the quintessence case, so it is interesting that these situations may be observationally favoured at the present epoch. In the final section we draw some general conclusions. Two appendices clarify the dimensions of various quantities used in our analysis and the connection between relevant energy conditions.

It should be stressed that this paper is not intended to be a review of the more general problem of black holes in a cosmological background, although that is of great interest in its own right [19–24]. It is also much more narrowly focussed than the earlier review of self-similar solutions in general relativity by Carr and Coley [25].

## II. POSITIVE-PRESSURE FLUID CASE

The Zel'dovich-Novikov argument [1] is based on a Bondi [26] analysis, in which one neglects the cosmological expansion and assumes that the background fluid will be swallowed from within an accretion radius  $R_A = 2GM/c_s^2$ , where  $c_s = \sqrt{\gamma - 1}c$  is the sound-speed and  $M$  is the black hole mass. The accretion rate for the black hole is then

$$\frac{dM}{dt} = 4\pi\rho c_s R_A^2 = \frac{16\pi G^2 M^2 \rho}{c_s^3} \quad (2.1)$$

where  $\rho$  is the mass density at the accretion radius, which is taken to be the background cosmological density. In the radiation-dominated era, which was the case considered by Zel'dovich and Novikov,  $c_s = c/\sqrt{3}$  and so  $R_A$  is just three times the Schwarzschild radius of the black hole. Since the mass density in a Friedmann radiation-dominated model is  $\rho = 3/(32\pi Gt^2)$ , where  $t$  is the cosmological time from the big bang, Eq. (2.1) then gives

$$\frac{dM}{dt} = \frac{KM^2}{t^2}, \quad K = \frac{9\sqrt{3}G}{2c^3}. \quad (2.2)$$

Note that this equation cannot be precisely correct. As discussed later, the density near the black hole is not exactly the background value and the accretion radius is not exactly  $2GM/c_s^2$ , but these effects merely modify the effective value of  $K$ . Equation (2.2) can now be integrated to give

$$M = \frac{M_0}{1 - \frac{KM_0}{t_0} \left(1 - \frac{t_0}{t}\right)}, \quad (2.3)$$

where  $M_0$  is the mass of the PBH at its formation time  $t_0$  and  $K^{-1}t_0$  is roughly the mass within the cosmological horizon at that time. (Both the particle horizon and Hubble horizon mass are exactly  $c^3t/G$  in the radiation era, which is a factor  $9\sqrt{3}/2 \approx 8$  larger.) For  $M_0 \ll K^{-1}t_0$ , Eq. (2.3) implies negligible accretion. However, for  $M_0 = K^{-1}t_0$ , it predicts  $M = K^{-1}t$ , in which case the black hole always has a size comparable to the particle horizon. For  $M_0 > K^{-1}t_0$ , the mass diverges at a time

$$t_\infty = t_0 \left(1 - \frac{t_0}{KM_0}\right)^{-1} \quad (2.4)$$

(i.e. at a time comparable to  $t_0$  unless  $M_0$  is very close to  $K^{-1}t_0$ ). These three behaviours are illustrated by the upper frame in Fig. 1.

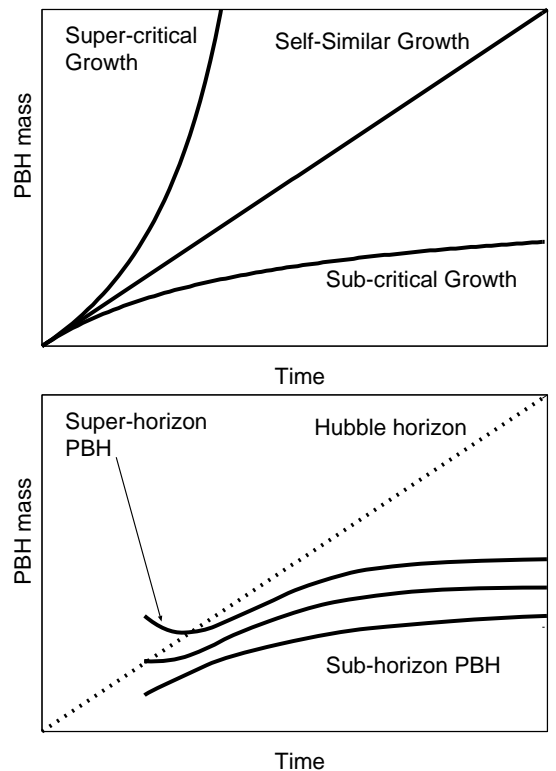


FIG. 1: Schematic figure showing the growth of PBHs with positive pressure for a Bondi-type analysis which neglects the cosmic expansion (top) and a full relativistic analysis which allows for it (bottom).

The middle case corresponds to self-similar growth and is precisely the situation which one might expect to pertain, since an overdense region needs to be larger than

the Jeans length at formation but not so large as to form a separate closed universe [27], both these length scales being of order the cosmological horizon size. This suggests that a PBH might in principle continue to grow as fast as the cosmological horizon until the end of the radiation-dominated era, when its mass would be of order  $10^{17}M_{\odot}$ . The existence of many such holes is excluded, for example, by observations of the large-scale bulk flows of galaxies [28], which might lead one to doubt that PBHs ever formed. Of course, the self-similar solution requires very fine-tuning of the black hole mass in the Zel'dovich-Novikov argument but, providing the black holes span a spectrum of masses, one might expect at least some of them to satisfy this condition.

That the behaviours shown the upper frame in Fig. 1 cannot be correct in a more precise calculation is indicated by the lower frame in Fig. 1. This shows the results of fully relativistic numerical calculations for the same initial black hole masses as in the upper frame [12]. We next investigate why the simple Bondi analysis fails and how it can be improved.

### A. Improved analysis neglecting cosmic expansion

The Zel'dovich-Novikov argument can be refined in various ways. In particular, Harada and Carr [29] improved the analysis by allowing the fluid to have a more general barotropic equation of state  $p = (\gamma - 1)\rho c^2$  with  $1 < \gamma < 2$  and by including a relativistic focussing factor  $\alpha$  in Eq. (2.1). In this case, the cosmological density is  $\rho = 1/(6\pi G\gamma^2 t^2)$ , so Eq. (2.3) remains valid but the factor  $K$  becomes

$$K = \frac{8G\alpha}{3c^3(\gamma - 1)^{3/2}\gamma^2}. \quad (2.5)$$

(Their expression was actually a factor of 4 smaller than this because they took the accretion radius to be  $GM/c_s^2$  rather than  $2GM/c_s^2$ .) The length-scale associated with the self-similar black hole solution is

$$R_S = \frac{2Gt}{c^2 K} = \frac{3c\gamma^2(\gamma - 1)^{3/2}t}{4\alpha} = \frac{c\gamma(\gamma - 1)^{3/2}}{2\alpha H} \quad (2.6)$$

(or a factor 4 larger in the Harada-Carr analysis), where  $H = 2/(3\gamma t)$  is the Hubble parameter. They also showed that an overdense region is a separate closed universe rather than part of our universe if its size exceeds

$$R_{\max} = \frac{2\sqrt{\pi}(3\gamma - 2)\Gamma\left(\frac{3\gamma - 1}{3\gamma - 2}\right)c}{3H\gamma\Gamma\left(\frac{3\gamma}{2(3\gamma - 2)}\right)}. \quad (2.7)$$

This expression applies for  $\gamma > 2/3$  and not just for  $1 < \gamma < 2$ . The self-similar scale is always less than the separate-universe scale in the present analysis but it goes above the self-similar scale for  $\gamma > 1.6$  in the original Harada-Carr analysis. They inferred that – even before attempting a more exact calculation – one should have reservations about the self-similar prediction for sufficiently hard equations of state.

Harada and Carr gave precise expressions for all relevant length scales in terms of the equation of state parameter and the functional forms are indicated in Fig. 2. The other scales shown there are the Jeans length [30] and the particle horizon size:

$$R_J = \frac{4\pi c\sqrt{\gamma - 1}}{(9\gamma - 4)H}, \quad R_P = \frac{2c}{(3\gamma - 2)H}. \quad (2.8)$$

Since there is some uncertainty in the appropriate accretion radius to use, various possibilities for the self-similar scale are shown; these all take the focussing factor  $\alpha$  to 1. Note that some of the scales in Fig. 2 can be extended down to  $\gamma = 2/3$  and they can all be extended above  $\gamma = 2$ .

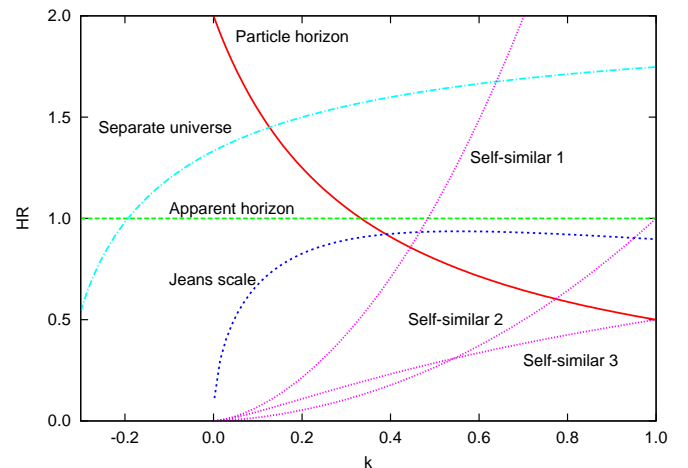


FIG. 2: The length scales associated with the Jeans condition, the particle horizon, the Hubble horizon, the separate-universe condition and the self-similar black hole solution are plotted against  $k \equiv \gamma - 1$ . Three self-similar scales are shown, corresponding to the accretion radius being (1)  $GM/c_s^2$ , (2)  $2GM/c_s^2$  and (3) the expression given by Eq. (2.12).

Babichev *et al.* [31] have further refined the accretion analysis by replacing Eq. (2.1) with

$$\frac{dM}{dt} = \frac{4\pi G^2 A M^2}{c^3} \left( \rho_{\infty} + \frac{p_{\infty}}{c^2} \right), \quad (2.9)$$

where the subscript  $\infty$  indicates asymptotic values and

$$A \equiv \frac{(3\gamma - 2)^{(3\gamma - 2)/2(\gamma - 1)}}{4(\gamma - 1)^{3/2}}. \quad (2.10)$$

This allows for the fact that the density of the fluid at the accretion radius is not the asymptotic cosmological value and it also incorporates the relativistic correction to Eq. (2.2) associated with the pressure. The Babichev *et al.* analysis is not confined to the  $p = (\gamma - 1)\rho c^2$  equation of state, but we assume this here because of the self-similarity requirement. Thus we again obtain Eq. (2.2) but with  $K$  is replaced by

$$K = \frac{2A(\gamma)G}{3\gamma c^3}, \quad (2.11)$$

so the mass for self-similar growth now has a more complicated dependence on the equation of state parameter than before. As  $\gamma$  decreases from 2 to 1, Eq. (2.10) indicates that the factor  $A$  increases from 4 to  $\infty$ , passing through  $6\sqrt{3}$  at  $\gamma = 4/3$  and diverging like  $[e/(\gamma - 1)]^{3/2}/4$  as  $\gamma \rightarrow 1$ . Comparison with Eq. (2.1) also shows that the “effective” accretion radius is

$$R_A = \frac{GM}{c^2} \sqrt{\frac{A\gamma}{\sqrt{\gamma-1}}} = \frac{2GM}{c_s^2} \sqrt{\left(\frac{A}{4}\right) \gamma(\gamma-1)^{3/2}}. \quad (2.12)$$

The last expression shows the deviation from the naive expression  $R_A = 2GM/c_s^2$ . For example, this is a factor  $\sqrt{2}$  for a stiff fluid and  $\sqrt{2/3}$  for a radiation fluid.  $R_A$  diverges in the limit  $\gamma \rightarrow 1$  and has the form

$$R_A \approx \frac{e^{3/4} GM}{2c_s^2}, \quad (2.13)$$

the numerical coefficient being almost exactly 1. The length-scale associated with the self-similar solution is now

$$R_S = \frac{2c}{AH} = \frac{8c(\gamma-1)^{3/2}}{(3\gamma-2)^{(3\gamma-2)/2(\gamma-1)} H} \quad (2.14)$$

and this is also shown in Fig. 2. Note that all three expressions for  $R_S$  shown there scale as  $(\gamma-1)^{3/2}$  as  $\gamma \rightarrow 1$  and the Babichev *et al.* expression coincides with the original Harada-Carr one in this limit. The ratio  $HR_S/c$  is 0.5 at  $\gamma = 2$  and asymptotes to  $8/(3\sqrt{3}) \approx 1.6$  as  $\gamma \rightarrow \infty$ , so the self-similar scale is always smaller than the separate-universe scale.

Clearly the above analysis does not apply for  $\gamma < 1$  since the value of  $A$  given by Eq. (2.10) is not well defined then. However, Babichev *et al.* argue that one can take  $A = 4$  in this situation, corresponding to the stiff fluid value. In this case, Eq. (2.9) gives a sensible accretion rate, even though the accretion radius given by Eq. (2.12) is not well defined. We will use this result in the later discussion.

## B. Improved analysis allowing for cosmic expansion

In all the above analyses, the existence of the self-similar black hole solution relates to the fact that the density scales as  $\rho \propto t^{-2}$  in a flat Friedmann background. However, none of the refinements discussed above allows for the cosmological expansion itself, which one might expect to oppose accretion, so one should be wary of the self-similar prediction.

Carr and Hawking’s refutation of the Zel’dovich-Novikov argument was based on the study of spherically symmetric self-similar solutions to Einstein’s equations [3]. Such solutions have the characteristic that all dimensionless quantities depend only on the dimensionless “similarity” variable  $z \equiv r/(ct)$ , where  $r$  is the comoving radial coordinate and  $t$  is the cosmological time [32]. The metric takes the form

$$ds^2 = -c^2 e^{2\Phi(z)} dt^2 + e^{2\Psi(z)} dr^2 + r^2 S(z)^2 d\Omega^2, \quad (2.15)$$

where  $d\Omega^2 \equiv d\theta^2 + \sin^2\theta d\varphi^2$ . A key role is played by the dimensionless “velocity” function

$$V(z) = |z| e^{\Psi-\Phi}, \quad (2.16)$$

which represents the fluid velocity in units of the speed of light relative to the similarity surfaces of constant  $z$  (which are expanding or collapsing spheres). Providing a certain regularity condition is satisfied [13], a value of  $z$  where  $V = 1$  corresponds to a null surface and is termed a “similarity horizon”. The flat Friedmann solution is itself self-similar and has

$$V(z) \propto z^{(3\gamma-2)/(3\gamma)}. \quad (2.17)$$

The point where  $V = 1$  corresponds to the cosmological particle horizon for  $\gamma > 2/3$ . (It corresponds to a cosmological event horizon for  $\gamma < 2/3$  but we assume  $1 < \gamma < 2$  in this section.) If one had a black hole in a Friedmann background, there would be two values of  $z$  for which  $V = 1$ , the inner one being associated with the black hole event horizon and the outer one with the cosmological particle horizon.

Values of  $z$  for which  $V = \sqrt{dp/d\rho}/c = \sqrt{\gamma-1}$  are also important because this may correspond to a sonic point (i.e. the pressure gradient may be discontinuous). For the radiation-dominated universe considered by Carr and Hawking, the sound speed is  $c/\sqrt{3}$ , so  $V = 1/\sqrt{3}$ . Since the Friedmann solution itself has such a point, one might envisage a self-similar solution in which one attaches a black hole *interior* to an exact Friedmann *exterior* via a sound-wave. If such a solution existed, it would represent a black hole growing at the same rate as the universe and  $V(z)$  would then have the form indicated by the broken curve in Fig. 3. However, Carr and Hawking showed that such a solution is not possible [3]. This implies that a black hole formed by local processes can never grow as fast as the universe, even if it starts off with a size comparable to the particle horizon. Thus the Zel’dovich-Novikov prediction cannot be correct in this case.

This conclusion was soon supported by detailed hydrodynamical calculations. If the formation and evolution of a PBH in a positive-pressure fluid with a local perturbation is simulated numerically, it is found that the PBH soon becomes much smaller than the cosmological horizon [33–35]. In this situation, one *would* expect the Zel’dovich-Novikov argument to apply, so one does not anticipate the black hole growing very much at all. This also applies for holes which are initially larger than the horizon [12], even though their mass should quickly diverge according to the Bondi-type analysis.

On the other hand, Carr and Hawking also constructed a one-parameter family of spherically symmetric self-similar PBH solutions for a radiation ( $\gamma = 4/3$ ) or dust ( $\gamma = 1$ ) fluid which they described as *asymptotically* Friedmann [3]. In these solutions, as  $z$  decreases, the velocity  $V$  decreases to a minimum between 1 and the sound-speed  $1/\sqrt{3}$ .  $V$  then rises through  $c$  at the black hole event horizon to infinity at the black hole singularity (corresponding to a finite value of  $z$  but zero physical distance). Two such solutions are shown in Fig. 3; we

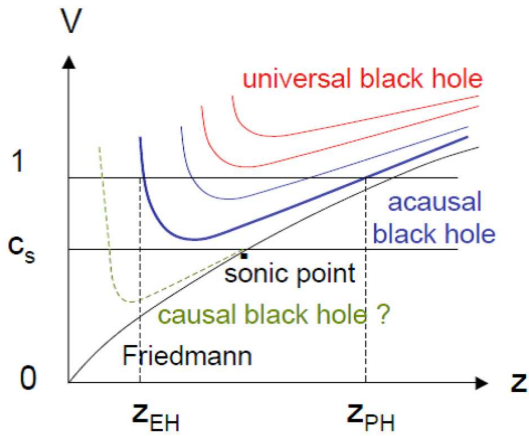


FIG. 3: Schematic figure of the function  $V(z)$  for self-similar asymptotically quasi-Friedmann black hole solutions, showing the two similarity horizons. The solutions are described by a single parameter and the minimum of  $V$  reaches the sonic value for solution. A black hole connected to an exact Friedmann background via a sound-wave would look like the dotted curve but such solutions cannot exist.

describe these as “acausal” since they are supersonic everywhere and the perturbation extends beyond the cosmological horizon. There are also solutions in which the minimum of  $V$  exceeds  $c$ , two examples of these being included. This corresponds to a “universal” black hole [36] in which there is no black hole or cosmological particle horizon. In this case, the entire universe might be regarded as being inside the black hole, although there are still black hole and cosmological apparent horizons [5].

The conclusion that there are self-similar black hole solutions in a universe which is asymptotically Friedmann but no such solutions in an exactly Friedmann universe was soon generalized to perfect fluids with equation of state  $p = (\gamma - 1)\rho c^2$  with  $1 \leq \gamma < 2$  [5, 6]. It is interesting to examine these asymptotically Friedmann solutions in more detail. The single parameter  $B_0$  which describes them is related to the perturbation at large  $z$  of the factor  $S(z)$  which appears in metric (2.15):

$$S = S_F(z)e^{B(z)}, \quad B \approx B_0 + B_1 z^{-2(3\gamma-2)/(3\gamma)}, \quad (2.18)$$

where  $B_1$  is determined by  $B_0$  and the value of  $\gamma$ . This is also related to the density perturbation at large distances:

$$W \equiv 8\pi G \rho r^2 / c^2 = W_F(z)e^{A(z)}, \quad A \approx A_1 z^{-2(3\gamma-2)/(3\gamma)}, \quad (2.19)$$

where  $A_1 = (2 - 3\gamma)B_1$ . For a sufficiently negative value of  $B_0$ , corresponding to a sufficiently large overdensity,  $V$  will reach a minimum as  $z$  decreases. This minimum will decrease as  $B_0$  increases but it will exceed 1 for  $B_0 < \beta_1 < 0$ , where the value  $\beta_1$  depends on  $\gamma$ . These solutions correspond to universal black holes.

On the other hand, the minimum will lie between  $V = 1$  and  $\sqrt{\gamma - 1}$  for  $\beta_1 < B_0 < \beta_2 < 0$ , corresponding to the black hole solutions discovered by Carr and Hawking. Here  $\beta_2$  also depends on  $\gamma$  and is the value of  $B_0$  for which

the minimum of  $V$  reaches  $\sqrt{\gamma - 1}$  [7]. Both these types of solution are also illustrated in Fig. 3. Note that the ratio of the black hole event horizon size to the particle horizon size decreases as  $B_0$  increases and so must always exceed the value associated with the solution with  $V_{\min} = \sqrt{\gamma - 1}$ . It might seem plausible to identify this with the Jeans criterion for gravitational collapse, in which case the ratio of the black event horizon and particle horizon should also have a minimum of order  $\sqrt{\gamma - 1}$ .

As  $B_0$  increases above  $\beta_2$ , the solution changes its form. Instead of  $V$  having a minimum, the solution reaches a sonic point, where the equations do not determine the behaviour uniquely, so there can be a discontinuity [7]. In fact, only a subset of solutions are “regular” at a sonic point in the sense that the pressure gradient is finite and they can be extended beyond there. There are two possible values of the pressure gradient, one associated with an isolated solution and the other with a 1-parameter family of solutions [37]. These regular transonic solutions do not contain black holes. Indeed, one can prove [3] that  $V$  cannot have a subsonic minimum and then rise through another sonic point to reach  $V = 1$  again. Instead, they pass smoothly to the origin and represent either overdense perturbations (for  $B_0 < 0$ ) or underdense perturbations (for  $B_0 > 0$ ) which grow at the same rate as the universe [7]. One can show that there is a continuum of underdense solutions and discrete bands of overdense ones [7, 37]. The solution with  $B_0 = 0$  is exactly Friedmann. These perturbed transonic Friedmann solutions are of great interest in their own right [38] but they are not relevant to the present considerations.

Since there is a one-parameter family of these solutions, we do not have to fine-tune the mass of the black hole to get self-similar growth as precisely as implied by Eq. (2.3). This is not surprising since the earlier analysis assumed that the accretion was driven by a local sound-wave rather than by an initial perturbation extending to infinity. However, the range of the parameter  $B_0$  for which self-similar growth is possible is still narrow and this reflects the proximity of the Jean length and separate universe conditions.

For comparison with the black hole solutions discussed later, Fig. 4 shows the Penrose diagrams for the acausal and universal black hole solutions. These correspond to Fig. 5 and Fig. 6 in Ref. [36], respectively. In these (and subsequent) Penrose diagrams a zig-zag line corresponds to a singularity, while a double-line corresponds to infinity. In the acausal case, it is worth emphasizing that the conformal structure is very different from that for a non-self-similar black hole in a Friedmann background. In particular, the black hole singularity is connected to the big bang singularity and necessarily naked for a while. In the universal case, the conformal diagram has two spacelike singularities and there are no null infinities. This contrasts with the diagram for a recollapsing Friedmann universe, which has two spacelike separated singularities.

Later it was realized that these self-similar black hole solutions are not “properly” asymptotic to Friedmann after all [8]. The solid angle at infinity is no longer  $4\pi$  because the radial area – measured by the quantity  $S(z)$

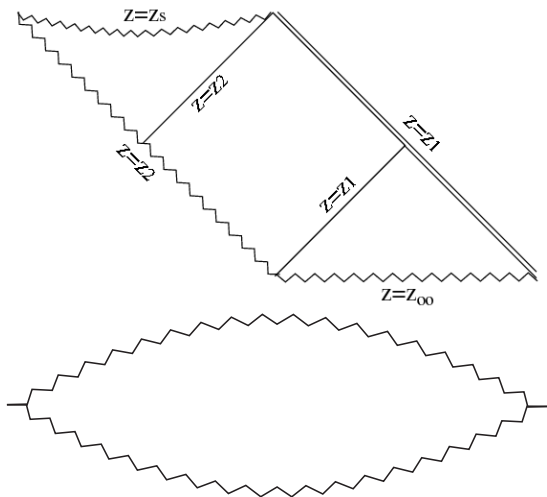


FIG. 4: The Penrose diagrams for an acausal black hole (top) and a universal black hole full (bottom) in an asymptotically quasi-Friedmann background.

– is perturbed out to arbitrarily large distances. More precisely, Eqs. (2.15) and (2.18) imply that the solid angle at infinity is  $4\pi e^{2B_0}$ , so there is a solid angle *deficit* in the black hole case ( $B_0 < 0$ ). These solutions may still be relevant to the real universe (since observations do not preclude such a deficit if it is not too large) but it would be more accurate to describe them as asymptotically “quasi-Friedmann”. Note that one tends to the separate-closed-universe condition as  $B_0 \rightarrow -\infty$  since the solid angle goes to zero. This is similar to the situation with the Barriola-Vilenkin monopole, which is an approximate general relativistic solution for the triplet Higgs scalar field in the region far from the center [39]. The spacetime is asymptotically flat in the monopole case but also contains a solid angle deficit. The prospects of detecting such a deficit by gravitational lensing observations have been reviewed in Ref. [40].

The complete classification of single-fluid shock-free positive-pressure spherically symmetric self-similar solutions by Carr and Coley [41] and Goliath et al. [42] has further clarified the situation. A key step in the Carr-Coley analysis is the derivation of all possible asymptotic behaviors at large and small distances from the origin. The behaviours at large spatial distances (usually corresponding to the limit  $z \rightarrow \infty$ ) are of four kinds: (1) asymptotically quasi-Friedmann (1-parameter); (2) asymptotically quasi-static (2-parameter). There are also two families of solutions which exist only when  $\gamma > 6/5$ : (3) asymptotically Minkowski at infinite  $z$  (1-parameter); and (4) asymptotically Minkowski at finite  $z$  but infinite physical distance (2-parameter). At small spatial distances, the solutions are also of four kinds: they contain either (a) a black hole or (b) a naked singularity at finite  $z$ ; or they can be connected to the origin at  $z = 0$  via a sonic point, in which case they are either (c) static or (d) a perturbation of the Friedmann solution there.

The complete family of  $1 < \gamma < 2$  solutions is found by combining the four kinds of large-distance behaviour

and four kinds of small-distance behaviour. The way in which one connects the large-distance and small-distance solutions depends crucially on whether or not there is a sonic point. If the solutions remain supersonic everywhere or subsonic everywhere, then the small- $z$  behaviour is uniquely determined by the large- $z$  behaviour. The asymptotically quasi-Friedmann black hole solutions are everywhere supersonic and of type (1a). However, if there is a sonic point, the solutions are more complicated because there can be a discontinuity. Indeed, we have seen that only a subset of solutions are regular at a sonic point. Although we focus here on the black hole asymptotically quasi-Friedmann solutions, it should be stressed that these represent only a small subset of the total family of self-similar solutions.

The study of critical phenomena in a Friedmann background [43–47] may also be relevant. If one considers the evolution of primordial density perturbations with a variety of forms, whose amplitude is described by some parameter  $p$ , one finds that a black hole forms if  $p$  exceeds some critical value  $p_*$  and that the black hole mass scales as

$$M \propto (p - p_*)^\beta \quad (2.20)$$

for  $p \approx p_*$ . The exponent  $\beta$  depends on the equation of state but is independent of the initial configuration of the collapsing fluid. The critical solution itself is self-similar and has been identified explicitly for different values of  $\gamma$  [48]. In terms of the Carr-Coley classification, it is either type (2d) (asymptotically quasi-static) for  $1 < \gamma < 1.28$  or type (3d) or (4d) (asymptotically Minkowski) for  $1.28 < \gamma < 2$ . (All the critical solutions have a regular centre and need to be asymptotically Friedmann there). Although the critical solution is only attained in the limit for which the black hole mass goes to zero, the critical solution is only an intermediate attractor, so there might still be a true self-similar attractor in which the mass is non-zero. In any case, we have seen that there are at least asymptotically quasi-Friedmann black hole solutions with non-zero mass for this range of  $\gamma$ .

### III. STIFF FLUID CASE

For the limiting case in which  $\gamma = 2$ , corresponding to a stiff fluid, the Zel’dovich-Novikov argument might conceivably apply because the sonic surface is a similarity horizon. Certainly this situation requires special consideration. Lin *et al.* at first claimed that there *is* a self-similar black hole solution in this case [9]. More precisely, they found a solution which has the exact Friedmann form outside the particle horizon (also a sonic point) but in which the velocity  $V$  inside this point has a minimum and then rises to 1 again, corresponding to a black hole event horizon. Such a solution resembles the ones shown in Fig. 3, except that it is subsonic between the event horizon and particle horizon.  $V$  therefore has a subsonic minimum, which is precisely the situation Carr and Hawking excluded in the  $1 < \gamma < 2$  case. The pressure



gradient appears to diverge at the event horizon but Lin *et al.* argued that this divergence can be removed by introducing an Eddington-Finkelstein-type coordinate.

Later Bicknell and Henriksen [10] showed that this solution is invalid because the second point where  $V$  reaches 1 as  $z$  decreases is not really a black hole event horizon. The reason for this is clarified in Ref. [13]. The metric induced on a constant- $z$  surface can be written as

$$ds^2 = -c^2 e^{2\Phi(z)} (1 - V(z)^2) dt^2 + r^2 S(z)^2 d\Omega^2, \quad (3.1)$$

which is why  $V = 1$  is usually a null surface. However, although Lin *et al.* pointed out that the pressure gradient diverges where  $V = 1$ , they failed to appreciate that the density goes to zero there and this implies that  $e^{\Phi(z)}$  diverges. This means that the time component of the induced metric is non-zero, so the surface  $V = 1$  is timelike rather than null and cannot be a black hole event horizon. However, it is worth stressing that Lin *et al.* correctly deduced the form of the solution up to the second point where  $V = 1$ .

Bicknell and Henriksen did manage to construct a numerical self-similar stiff solution containing a PBH attached to an exact Friedmann exterior [10]. However, this solution has the feature that the stiff fluid turns into null dust at the timelike surface where  $V = 1$ , with the fluid becoming lightlike there. This situation might appear to be contrived but Reed and Henriksen [49] were able to find a solution of this kind by generalizing a previous solution of Hacyan [50] for a radiation-dominated model to arbitrary  $\gamma$ . In the Hacyan solution, the black hole is formed by inflowing radiation in a region described by the Vaidya metric. Reed and Henriksen showed that this solution is necessarily self-similar, with the black hole mass growing asymptotically as  $t$ , but it is generally unphysical because it involves a decompression wave (which they claim violates the 2nd law of thermodynamics). However, they argued that it could be physical in the stiff case and that this might allow PBHs to form with masses up to  $4M_\odot$  if the equation of state is stiff up to the end of the hadron era. Although this seemed plausible at the time, the modern view does not favour a stiff hadron era, so this solution now also appears contrived.

Recently, Harada *et al.* [13] have studied the stiff self-similar solutions with an exact Friedmann exterior in greater depth by treating the fluid equations as a 2-dimensional autonomous system and exploiting the equivalence (discussed below) between a stiff fluid and a massless scalar field. They find a family of extensions within the inner  $V = 1$  point which contain a massless scalar field. These might appear more natural than the Bicknell-Henriksen or Reed-Henriksen solutions but they still enter a negative mass region, which is presumably unphysical.

Harada *et al.* also prove that there is no asymptotically Friedmann or asymptotically quasi-Friedmann solution with a sonic point containing a black hole in the stiff case. This is analogous to the situation in the non-stiff case, where solutions with sound-waves can only represent density perturbations at the origin rather than black holes. The possibility of self-similar solutions is even more re-

stricted in the stiff case because there is no range for the minimum of  $V$  between 1 and the sonic value. However, as discussed below, one might still expect the existence of universal self-similar black hole solutions (with  $V_{min} > 1$ ) in the stiff case, these being directly analogous to the supersonic solutions discussed in Section II.B.

#### IV. SCALAR FIELD CASE

Scalar fields are a key ingredient in modern cosmology. They play an important role during inflation and certain phase transitions; they feature in the preheating scenarios; and they are pervasive in all sorts of alternative theories (e.g. string theory and scalar-tensor theory) which are likely to be relevant in the high curvature phase of the early universe. In such scenarios, scalar fields necessarily dominate the energy density of the universe at some stage of its evolution. One important physical feature in this case is that the scalar field can contain sound-waves, with the sound-speed ( $\sqrt{\delta p / \delta \rho} \neq \sqrt{p / \rho}$ ) being the speed of light in the short-wavelength limit [29]. Equation (2.12) then suggests that the accretion radius is  $\sqrt{2}$  times the Schwarzschild radius. As discussed later, this is different from the case of a perfect fluid with negative pressure but constant  $\gamma$ , in which there are no sound-waves.

##### A. Massless scalar field

If there is no scalar potential, the scalar field is massless and one might expect the same results as in the stiff fluid analysis, since a scalar field is equivalent to a stiff fluid provided the gradient is everywhere timelike [11]. This is also consistent with numerical simulations of PBH growth in a Friedmann universe with a massless scalar field [12]. In fact, Harada *et al.* [13] have constructed an analytic proof that there is no self-similar PBH solution in an exact or asymptotically Friedmann background if the universe is dominated by a massless scalar field. The proof exploits Brady's formulation of the self-similar equations as an autonomous system [51] and uses Bondi coordinates  $(v, r, \theta, \varphi)$  and the variables

$$y(\xi) \equiv 1 - \frac{2GM}{c^2 r}, \quad z(\xi) \equiv \frac{r}{vg}. \quad (4.1)$$

Here  $v$  is the ingoing radial null coordinate and the variable  $z$  is to be distinguished from the quantity used elsewhere to denote the similarity variable.  $M$  is the Misner-Sharp mass,  $\xi \equiv \ln(r/|cv|)$  and  $g \equiv g_{vv}/g_{vr}$ .

The trajectories of the Harada *et al.* solutions in the  $(y, z)$  plane are illustrated in Fig. 5. The shaded regions are prohibited for a real scalar field. The gradient of the scalar field can change from timelike to spacelike on  $\gamma = 0$  (thin solid line) or  $2z(\gamma + \kappa) = \gamma$  (thin dashed line) or  $y = 0$ , where

$$\gamma \equiv -\kappa \pm \sqrt{\frac{1 + 4\pi\kappa^2 - y^{-1}}{4\pi(1 - 2z)}}. \quad (4.2)$$

Here  $\kappa$  is a constant, which is taken to be  $1/\sqrt{12\pi}$ . Two numerical solutions within the particle horizon are shown: the one above the Friedmann line is the positive branch and goes directly to the negative mass region ( $y > 1$ ), while the one below Friedmann reaches  $y = 3/4$  before going to the negative mass region. The Bicknell-Henriksen null-dust black hole solution connects to the lower numerical solution, but this is never realized by a scalar field. The black hole has an event horizon where the solution intersects  $z = 1/2$  and an apparent horizon at  $y = 0, z = \infty$ .

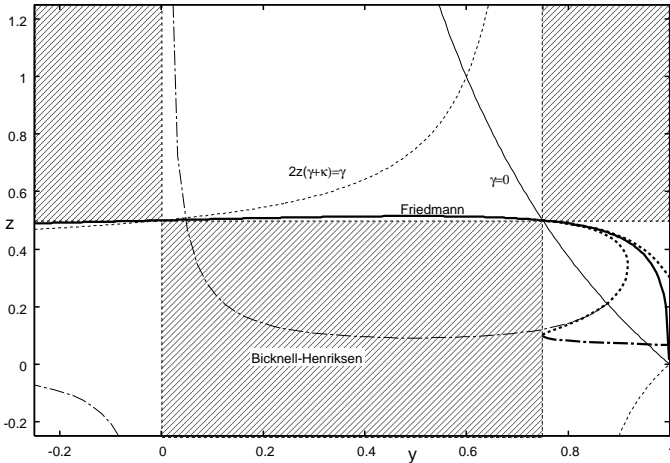


FIG. 5: The massless scalar field solutions in the  $(y, z)$  plane. The expanding flat Friedmann solution is plotted as a thick solid line. Two numerical solutions within the particle horizon, whose significance is described in the text are shown by the thick dotted lines. The Bicknell-Henriksen null-dust black hole solution is indicated by a thin dashed-dotted line.

Recent work by Bhattacharya *et al.* [52] has found explicit self-similar asymptotically Friedmann black hole solutions with  $V_{min} > 1$  in the massless scalar field case. These correspond to universal black hole solutions [36] and are a special subset of the Roberts self-similar solutions [53], the latter describing the inhomogeneous spherically symmetric gravitational collapse of a massless scalar field which is minimally coupled to gravity (see also Ref. [54]). The corresponding solution in the stiff case was obtained by Gutman and Bespal'ko [55], the equivalence between these solutions being shown in Appendix A of Ref. [56].

## B. Massive scalar field

The situation is more complicated – but also more interesting – if there is a scalar potential (i.e. if the scalar field has mass, this being associated with second derivative of the potential). In particular, this applies in the quintessence scenario, in which the scalar field rolls down a flat potential [57]. In this context, Bean and Magueijo [16] have used a variant of the Zel'dovich-Novikov argument to claim that PBHs could grow up to  $\sim 100M_{\odot}$  through the accretion of quintessence before

nucleosynthesis and this could be large enough to provide the seeds for the supermassive black holes found in galactic nuclei. A generalization of this analysis by Custodio and Horvath [58], though not involving self-similar growth, has also claimed there could be appreciable accretion of quintessence in some circumstances.

We now examine these arguments in more detail. For clarity the physical dimensions of various quantities appearing in this section are justified in Appendix A. Bean and Magueijo [16] consider a scalar field with a potential of the form

$$V(\phi) = V_0 e^{-\sqrt{8\pi G} \lambda \phi/c^2} \quad (4.3)$$

where  $\lambda$  is a dimensionless constant. This is the only form of the potential compatible with a self-similar solution, so one can immediately exclude the possibility of a black hole growing as fast as the universe with any other form. Note that  $\phi$  has units  $c^2/\sqrt{G}$  and  $V$  has units  $c^4/(GL^2)$  where  $L$  is a lengthscale.

By generalizing an analysis of Jacobson [59], which attaches a Schwarzschild solution to a cosmological background in which the field has the asymptotic value  $\phi_c(t)$ , they claim that the energy flux through the event horizon is  $\dot{\phi}_c^2/c$ , where a dot denotes  $d/dt$  and  $t$  is cosmic time. This leads to a black hole accretion rate [60]

$$\frac{dM}{dt} = 16\pi G^2 M^2 \dot{\phi}_c^2 / c^7. \quad (4.4)$$

They correctly note that only the kinetic energy of the scalar field is accreted. The scalar potential merely influences the evolution of the background field, the potential and the cosmic scale factor:

$$\phi_c = -\frac{2c^2}{\lambda\sqrt{8\pi G}} \ln\left(\frac{t_1}{t}\right), V = V_0 \left(\frac{t}{t_1}\right)^{-2}, a = \left(\frac{t}{t_1}\right)^{2/\lambda^2} \quad (4.5)$$

where  $t_1$  is defined by

$$t_1^2 \equiv \frac{c^2(6 - \lambda^2)}{4\pi G V_0 \lambda^4}. \quad (4.6)$$

The expansion is accelerating providing  $\lambda < \sqrt{2}$ . Since a scalar field has

$$\rho c^2 = \frac{1}{2c^2} \dot{\phi}^2 + V, \quad p = \frac{1}{2c^2} \dot{\phi}^2 - V, \quad (4.7)$$

Eq. (4.4) can also be inferred from Eq. (2.9) of the Babichev *et al.* analysis if one assumes  $A = 4$ . From Eq. (4.4) and the expression for  $\phi_c$  given by Eq. (4.5), the accretion rate now becomes

$$\frac{dM}{dt} = \frac{KM^2}{t^2}, \quad K = \frac{8G}{c^3 \lambda^2}. \quad (4.8)$$

This equation can be integrated to give Eq. (2.3) but with the revised expression for  $K$ .

Note that Akhoury *et al.* [61] have recently claimed that the standard Bondi accretion formula is not applicable for a quintessence fluid. Their argument is not specific to black holes but just assumes the existence of

a Killing vector and the invariance of the stress-energy tensor. For a single canonical scalar field, they show that the self-similarity assumption restricts the form of the potential to be exponential. However, we have already inferred this result above and their argument does not seem to invalidate our specific accretion analysis. We therefore assume that Eq. (4.8) still applies.

As before, one has negligible growth for  $M_0 \ll K^{-1}t_0$ , self-similar growth for  $M_0 \sim K^{-1}t_0$  and rapid divergence for  $M_0 \gg K^{-1}t_0$ . However, the self-similar mass is now of order the mass within the Hubble horizon (which is also the cosmological apparent horizon), since the particle horizon does not exist if the acceleration extends back to the infinite past. More precisely, using the scale factor given by Eq. (4.5), the Hubble mass can be shown to be

$$M_H = \frac{3c^3\lambda^2 t}{16\pi G} = \frac{3t}{2K}, \quad (4.9)$$

which is just a factor 3/2 larger than the mass corresponding to the self-similar solution. The self-similar result is the one which Bean and Magueijo exploit. However, their analysis is exactly equivalent to that of Zel'dovich and Novikov [1], which Carr and Hawking proved was inapplicable. It is therefore important to know whether the non-existence of a black hole self-similar solution extends to this case.

Custodio and Horvath [58] also consider black hole accretion of a quintessence field with a scalar potential but, instead of Eq.(4.4), they use

$$\frac{dM}{dt} = 27\pi G^2 M^2 \dot{\phi}_c^2 / c^7, \quad (4.10)$$

where the numerical factor differs because of relativistic beaming. They disagree with Bean and Magueijo's choice of the function  $\phi_c$  in Eq. (4.5) on the grounds that it neglects the local decrease in the background scalar field resulting from the accretion. For reasons which are not altogether clear, they therefore focus on a model in which the quintessence flux into the black hole is constant. This requires that the potential have a specific form, which is different from (4.3) and therefore incompatible with self-similarity. This leads to the mass evolution

$$M = \frac{M_0}{1 - \frac{K'M_0}{t_0}(t - t_0)}, \quad (4.11)$$

where  $K'$  is a constant related to the (fixed) flux. This is equivalent to the Zel'dovich-Novikov formula (2.3) providing one replaces  $K$  by  $K't$ . In particular, the mass diverges at a time

$$t_\infty \equiv t_0 \left( 1 + \frac{1}{K'M_0} \right) \quad (4.12)$$

and they attribute this unphysical feature to the fact that the constant-flux assumption must fail. They therefore consider alternative models in which the flux decreases as a power of time. In general, they find that the increase in the black hole mass is small unless its initial value is finely tuned.

Although Custodio and Horvath conclude that the Bean-Magueijo argument is wrong, we would claim that they have not identified the more significant problem – that it neglects the background cosmological expansion. Indeed, since the cosmological expansion is neglected in both these analyses, they could both be flawed for the same reason as the original Zel'dovich-Novikov one.

To examine the relativistic situation, we seek a spherically symmetric self-similar solution which represents a black hole in a quintessence-dominated cosmological background. In the quintessence scenario, for which the similarity assumption requires the potential to have the form given by Eq. (4.3), the analysis in Ref. [13] shows that there is no self-similar black hole solution in an exact or asymptotically Friedmann or asymptotically quasi-Friedmann background if the universe is decelerating and no such solution in an exact Friedmann background if it is accelerating. However, this does not prove non-existence for the case in which the universe is accelerating and asymptotically Friedmann, which is what would be required to disprove the Bean-Magueijo proposal.

Kyo *et al.* [62] have shown that, with an accelerating scalar potential, there is a one-parameter family of self-similar solutions which are properly asymptotically Friedmann, in the sense that there is no solid angle deficit. In such solutions, the perturbation falls off very rapidly at spatial infinity. The question of whether such solutions can have a black hole event horizon remains open. However, as discussed in the next section, there is a self-similar black hole solution in a closely-related scenario.

## V. DARK ENERGY FLUIDS

The observed acceleration of the universe means that violation of the strong energy condition is required. The definition of the various energy conditions is given in Appendix B; for a perfect fluid with  $p = (\gamma - 1)\rho c^2$  and positive  $\rho$ , the violation of the strong one corresponds to  $\gamma < 2/3$  (including negative values). We describe such a fluid as “dark energy”. However, it should be stressed that this term is used in various ways by different authors and indeed is sometimes used interchangeably with the term “quintessence”, so some clarification of terminology is required. In general, a scalar field corresponds to a stiff fluid ( $\gamma = 2$ ) when the potential is negligible and to a cosmological constant ( $\gamma = 0$ ) when the potential dominates, so it corresponds to a fluid with  $0 < \gamma < 2$  at intermediate times and to “dark energy” (in our sense of the term) for any period when  $0 < \gamma < 2/3$ . Quintessence corresponds to the special case in which the scalar field rolls down a very flat potential, so that  $\gamma$  may change but is always in the latter range [57]. We have seen that there are sound-waves in the scalar field case but there are no sound-waves for a perfect fluid with  $0 < \gamma < 1$  because  $\sqrt{p/\rho} = \sqrt{\gamma - 1}c$  is imaginary.

The question of whether a PBH can grow in a self-similar manner if the universe is dominated by dark energy can be addressed using the analysis of Babichev *et*

*al.* [63]. Although the parameter  $A$  given by Eq. (2.10) is not well-defined for  $0 < \gamma < 1$ , Babychev *et al.* argue that it can be taken to be 4 throughout this range. In this case, Eqs. (2.2) and (2.11) can still be applied, which suggests that self-similar growth is again possible. However, note that  $K \rightarrow \infty$  as  $\gamma \rightarrow 0$ , so the self-similar black hole has zero mass in this limit.

As in the positive-pressure case, one should be wary of the self-similar claim since the above analysis neglects the cosmic expansion. The exact relativistic analysis has been carried out in our two recent studies of self-similar solutions with dark energy [17, 18]. These use a combination of numerical and analytical methods to classify all spherically symmetric self-similar solutions with  $0 < \gamma < 2/3$  which are asymptotically Friedmann at large distances. This extends the Carr-Coley classification of self-similar solutions with  $1 < \gamma < 2$  to the  $0 < \gamma < 2/3$  case. In fact, Ref. [64] already includes an asymptotic analysis of such solutions, although there are some errors in that work.

There are several differences from the  $1 < \gamma \leq 2$  situation. First, the fact that there are no sound-waves changes – and indeed simplifies – the analysis, since there can be no discontinuities and solutions are analytic everywhere. Second, the limiting values of  $z$  for large and small spatial distances are reversed in the Friedmann solution: large spatial distances now correspond to  $z \rightarrow 0$  and small ones correspond to  $z \rightarrow \infty$ . Furthermore, we will find that some solutions necessarily span both positive and negative  $z$  values. There are also several differences as regards the solutions themselves. For example, one loses the exact static solution if  $0 < \gamma < 2/3$  because one needs positive pressure to balance gravity. This also means that there are no asymptotically static ones, although there are still what are termed asymptotically quasi-static solutions (a point missed in Ref. [64]). On the other hand, one gains some solutions, since the Kantowski-Sachs and asymptotically Kantowski-Sachs solutions [65] now become physical. Although one loses the solutions for  $6/5 < \gamma < 2$  which are asymptotically Minkowski at *large* distances [41], one gains solutions which are asymptotic to a negative-mass Schwarzschild singularity at *small* distances.

The  $0 < \gamma < 2/3$  case is particularly important in the cosmological context because there then exists a one-parameter family of self-similar solutions which are properly asymptotic to the flat Friedmann model at large distances. This has also been pointed out in Ref. [66]. As we have seen, this does not apply in the positive pressure case because the solutions are then only asymptotically quasi-Friedmann. However, these solutions are still acausal, in the sense that there is no sonic point, so they are not like the solutions which Carr and Hawking were originally seeking in the positive pressure context.

In order to investigate these solutions, we have used the asymptotic analysis of Ref. [17] and integrated the field equations numerically in Ref. [18]. We find that the solutions which are asymptotically Friedmann at large distances ( $z \rightarrow 0$ ) are described by a single parameter  $A_0$ , which is analogous to the parameter  $B_0$  in the  $1 < \gamma < 2$

case and may be interpreted as the density perturbation at spatial infinity. However, the expressions for  $A$  and  $B$  in Eqs (2.18) and (2.19) no longer apply and must be replaced by

$$A \approx A_0 z^{(2-\gamma)/\gamma}, \quad B \approx -(A_0/6\gamma) z^{(2-\gamma)/\gamma} \quad (5.1)$$

in the  $0 < \gamma < 2/3$  regime. The form of the solutions changes as one varies the parameter  $A_0$  and one has three possible behaviours [18]. The first class of solutions (with  $A_0$  positive) is confined to the  $z > 0$  domain and contains a naked singularity; the Friedmann solution itself has  $A_0 = 0$ . The second class (with  $A_0$  negative but not too small) involves an extension into the  $z < 0$  region and contains a black hole; this is analogous to the Carr-Hawking solution in the positive-pressure case. The third class (with  $A_0$  even more negative) also involves an extension into the  $z < 0$  region and corresponds to a wormhole solution; this has no analogy in the positive-pressure case.

The form of  $V$  in these solutions is shown as a function of  $1/z$  in Fig. 6 for the  $\gamma = 1/3$  case and we discuss them below in more detail. Note that  $V = 1$  now corresponds to a cosmological event horizon rather than a particle horizon. Although we did not investigate all values of  $\gamma$  in the range 0 to  $2/3$ , we would expect the results to be qualitatively similar. We note that the system of equations was also numerically investigated in Ref. [66] but no black hole or wormhole solutions were reported there because the analytic extension of the solutions to the  $z < 0$  domain was not considered.

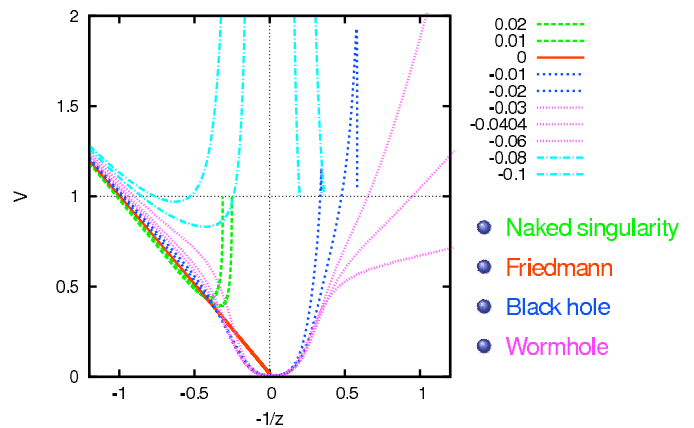


FIG. 6: The form of  $V$  in the negative pressure solutions with  $\gamma = 1/3$ , showing the transition from the black hole to naked singularity solutions as one varies the parameter  $A_0$  which describes their asymptotic behaviour. The ordinate is taken to be  $-1/z$ , which is large and negative in the asymptotic Friedmann region and small near the origin if the solution reaches there.

### A. Black hole solutions

Solutions with  $\alpha_1 \leq A_0 < 0$  for some critical  $\gamma$ -dependent value  $\alpha_1$  (i.e. which are not too overdense) are

asymptotically quasi-Kantowski-Sachs as  $z \rightarrow \infty$ . The exact Kantowski-Sachs metric has the form

$$ds^2 = -\frac{(2-3\gamma)(2-\gamma)}{\gamma^2}c^2 dt^2 + t^{4(1-\gamma)/\gamma} dr^2 + c^2 t^2 d\Omega^2 \quad (5.2)$$

with a suitable radial coordinate and this means that the physical distance tends to a finite limit (like a shell) rather than zero as  $z \rightarrow +\infty$ . However, these solutions can be extended analytically into the region with negative  $z$  and, as  $z$  increases from  $-\infty$ , they then approach a positive-mass singularity at some negative value  $z_*$ . Since the Kantowski-Sachs solution has a curvature singularity at  $t = 0$ , the analytic extension in this case must be interpreted as an extension from the positive  $r$  to negative  $r$  region. However, the areal radius and mass are still positive in the extended region. This solution describes a black hole in an asymptotically Friedmann background and is therefore analogous to the Carr-Hawking solution in the positive pressure case. In particular, there are two similarity horizons, corresponding to a cosmological event horizon at  $z = z_1$  and a black hole event horizon at  $z = z_2$ , and a black hole singularity at  $z = z_*$ . However, the initial big bang singularity is null rather than spacelike in this case. The Penrose diagram for this solution is shown in Fig. 7 and is reproduced from Fig.12 of Ref. [18].

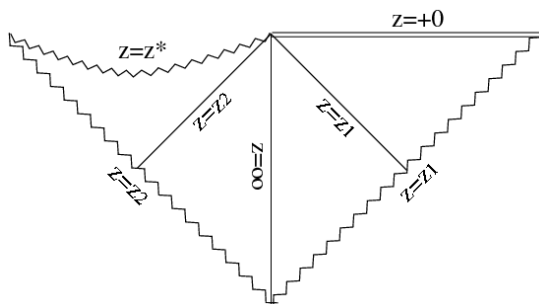


FIG. 7: The Penrose diagram for a solution with negative pressure which contains a black hole in an asymptotically Friedmann universe.

For  $\gamma = 1/3$ , we have found numerically that the ratio of the size of the black hole event horizon to the Hubble length goes from 0 and 0.70 as  $A_0$  decreases from 0 to  $\alpha_1$ . This means that the size of a self-similar cosmological black hole has an upper limit, so only sufficiently small black holes can grow as fast as the universe, but may be arbitrarily small. Although a black hole may not grow as fast as the universe when it first forms (if it is as large as the background Hubble length), the ratio of its radius relative to the Hubble length will decrease as time proceeds. Eventually it will fall below 0.70, after which the black hole starts to grow self-similarly.

Since there is a one-parameter family of these solutions, we do not have to fine-tune the mass of the black hole in the way implied by Eq. (2.3) to get self-similar growth. This also applied in the positive pressure case but the parameter range is now even more extended. Also

the positive-pressure solutions may be less plausible since they are only asymptotically quasi-Friedmann.

## B. Wormhole and white hole solutions

Solutions with  $\alpha_3 \leq A_0 < \alpha_1$  (where  $\alpha_3$  is another  $\gamma$ -dependent critical value) are also asymptotically quasi-Kantowski-Sachs as  $z \rightarrow \infty$ . These solutions can again be analytically continued into the negative  $z$  region but they now reach  $z = 0^-$  (corresponding to infinite physical distance) instead of a singularity. They are asymptotically Friedmann at  $z = 0^+$  but asymptotically quasi-Friedmann at  $z = 0^-$ , except for a special value  $\alpha_2$  between  $\alpha_1$  and  $\alpha_3$  which allows the solution to be asymptotically Friedmann at both ends. Since the physical radius never becomes zero except at the big bang singularity, these solutions correspond to cosmological wormholes and there are two cosmological event horizons at  $z_1$  and  $z_2$ . The Penrose diagram for these solutions is shown in Fig. 8 and is reproduced from Fig. 15 of Ref. [18]. This only differs from Fig. 7 in that there is no singularity at the top left.

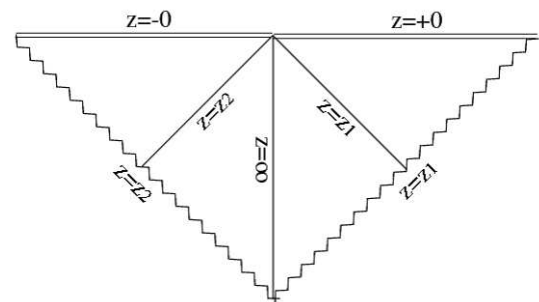


FIG. 8: The Penrose diagram for solutions with negative pressure in which a wormhole connects an asymptotically Friedmann solution on the right to an asymptotically quasi-Friedmann universe on the left.

The transition from the black hole to wormhole solutions has a simple physical explanation. As the parameter  $A_0$  becomes more negative, the ratio of the size of the black hole apparent horizon to the size of the cosmological apparent horizon increases. The transition occurs when this ratio reaches unity and thereafter both horizons disappear, every region being trapped. This is in contrast to the  $1 < \gamma < 2$  case, where the two apparent horizons never merge and one tends to a separate closed universe as the black hole size increases [5].

The solutions with  $A_0 < \alpha_3$  are asymptotically quasi-static rather than asymptotically quasi-Kantowski-Sachs as  $z \rightarrow \infty$  but they can still be extended into the negative  $z$  region. These solutions have a central curvature singularity at some negative value  $z_*$ , where the mass is positive, and correspond to the formation of a wormhole from a white hole in a Friedmann background at  $t = 0$ . The Penrose diagram for this solution is shown in Fig. 9 and is reproduced from Fig.15 of Ref. [18].

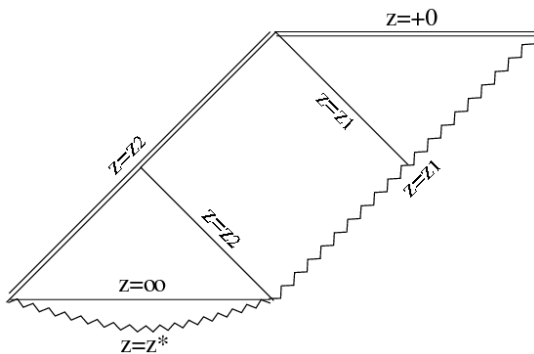


FIG. 9: The Penrose diagram for a solution with negative pressure which describes the formation of a wormhole from a white hole in an asymptotically Friedmann spacetime.

In order to understand these wormhole solutions, some clarification of what is meant by a wormhole is required. This is discussed in detail in another paper [67] but we summarize the main points here. A wormhole is an object in general relativity which connects two or more asymptotic regions. The most famous example of a static wormhole is the Morris-Thorne solution [68] and it is well known that this requires a violation of the null energy condition [69, 70]. The definition of this is given in Appendix B but it corresponds to  $\mu + p \geq 0$  for a perfect fluid. Dynamical wormholes are not so well understood but their study was pioneered by Hochberg and Visser [71] and Hayward [72], who defined a wormhole throat in a dynamical spacetime as some kind of trapping horizon [73]. In their definitions, a wormhole throat is a two-dimensional surface of non-vanishing minimal area on a null hypersurface and the null energy condition must again be violated there. However, there is no past null infinity in the context of our cosmological wormhole solutions because there is an initial singularity. In fact, the Hochberg-Visser and Hayward definitions are inapplicable because the spacetime is trapped everywhere and there is no trapping horizon. This demonstrates that these definitions miss the important class of *cosmological* wormholes which are asymptotically Friedmann and start with a big bang.

In order to remedy this problem, we define a wormhole throat as a two-dimensional surface of non-vanishing minimal area on a *spacelike* hypersurface [67]. The one-parameter family of spherically symmetric self-similar wormhole solutions in an accelerating Friedmann background discussed above satisfy this definition. They are asymptotically Friedmann at one infinity and they have another infinity, which may also be asymptotically Friedmann for a special value of the parameter. Interestingly, the dominant energy condition is satisfied everywhere.

In Ref. [67] we construct two analytic examples of self-similar cosmological wormholes. One corresponds to the numerical solution obtained above, but it contains a singular hypersurface which violates the null energy condition. The other is a smooth model involving a combination of a perfect fluid and a ghost scalar field (i.e. a

massless scalar field with a negative kinetic term) but the total matter content still satisfies the dominant energy condition.

## VI. EXOTIC CASES

In this section, we consider values of  $\gamma$  which are outside the “conventional” range from 0 to 2 but might nevertheless arise in more exotic scenarios. These cases have not been studied in sufficient detail to come to definite conclusions. In particular, the spherically symmetric self-similar solutions have not been obtained explicitly. Nevertheless, it is worthwhile summarizing what is known about these situations.

### A. Negative pressure fluids satisfying SEC

A fluid with  $2/3 < \gamma < 1$  has negative pressure but it is not sufficiently negative to violate the strong energy condition. Self-similar solutions in this context have been studied in Ref. [64], which shows that their possible large and small distance behaviours include asymptotically quasi-Friedmann and asymptotically Kasner (corresponding to a black hole singularity), respectively. Indeed, the equations are formally similar to those in the  $1 < \gamma < 2$  case, except for the absence of sound-waves and that should not be relevant if one is seeking an analogue of the  $1 < \gamma < 2$  solutions which are supersonic everywhere. One would therefore *expect* the black hole solutions to exist, although this has not yet been demonstrated numerically.

### B. Phantom fluids

A fluid with  $\gamma < 0$  is described as a “phantom”. More generally, this corresponds to any fluid with  $p < -\rho c^2$  and this possibility has been explored in the context of fluids with equation of state of the form  $p = \alpha c^2(\rho - \rho_0)$  in Ref. [74]. Here we confine attention to the situation with  $\alpha = \gamma - 1$  and  $\rho_0 = 0$  because of the similarity assumption. There are still asymptotically Friedmann self-similar solutions in this case, so the previous analysis can in principle be extended to cover this. However, since the scale factor is  $a(t) = |t/t_1|^{2/(3\gamma)}$ , there is a “big rip” singularity [75, 76] at  $t = 0$ , spacelike infinity is at  $z = +\infty$  and the domain of definition is  $t < 0$ . Note that there is no future null infinity in these solutions, so there is no black hole event horizon. However, there is still a black hole apparent horizon and one may take this to define the mass of the black hole. Equation (2.9) then has the consequence that the PBH mass should decrease as a result of phantom accretion and it would need to decrease in proportion to the time left to the big rip in a self-similar solution.

Babichev *et al.* [63] have analysed black hole accretion of phantom energy and find that the mass evolves

according to

$$M = \frac{M_i}{1 + \frac{M_i}{M_0 \tau} \frac{t'}{\tau - t'}} \rightarrow \dot{M}_0 |t|, \quad (6.1)$$

where  $M_i$  is the initial mass and  $\dot{M}_0 \equiv (3\gamma c^3/2GA)$  with  $A$  taken to be 4. The first expression uses the Babichev *et al.* time coordinate with the time of the big rip corresponding to  $t' = \tau$ ; the second expression applies in the limit  $t' \rightarrow \tau$  and uses our time coordinate  $t = t' - \tau$ . Note that Eq. (6.1) is very similar to Eq. (2.3) with  $K$  being given by Eq. (2.11). Although the mass evolution is clearly self-similar close to the big rip, the derivation of Eq. (6.1) again fails to account for the background expansion, since the black hole is not embedded in a cosmological background. Nevertheless, self-similar black hole solutions might in principle exist in this situation. The fact that black hole mass is *shrinking* rather than increasing is of obvious astrophysical interest. There have also been studies of black hole accretion in a ghost condensate [77, 78]. However, this term should not be confused with the word “ghost” which is sometimes used for the phantom fluid itself. A ghost condensate is more like dust ( $\gamma = 1$ ) than a phantom fluid.

### C. Tachyonic fluids

A fluid with  $\gamma > 2$  might be described as “tachyonic”, in the sense that the sound-speed exceeds  $c$ . It formally contradicts the dominant energy condition but not the other ones. In fact, most of the equations in the  $1 < \gamma < 2$  case be formally extended into this domain. In particular, all the curves in Fig. 2 can be extended to arbitrarily large  $\gamma$ . Babichev *et al.* [79, 80] have considered black hole accretion in this case but the problem does not yet seem to have been studied in the cosmological context.

## VII. DISCUSSION AND SUMMARY

In this review, we have discussed the possible existence of self-similar solutions containing a black hole or wormhole in an asymptotically Friedmann background whose density is dominated by a perfect fluid with  $p = (\gamma - 1)\rho c^2$  or a scalar field. A simple Bondi-type analysis predicts an accretion rate of the form given by Eq. (2.2) in all cases, though with a different value of the constant  $K$ , which appears to permit self-similar growth. Ultimately, this is because the Friedmann equation implies that the density scales as  $\rho \propto t^{-2}$  in a flat Friedmann background, which is precisely the condition for self-similarity.

However, the simple analysis is suspect because it neglects the cosmic expansion, so this has motivated a more careful relativistic analysis which allows for the expansion. In the positive-pressure case, the only known self-similar black hole solution with a sound-wave which is *exactly* Friedmann at large distances is the Bicknell-Henriksen stiff-fluid one; but even this is contrived since

it requires the fluid to turn into null dust. There is also the Hacyan radiation-dominated solution (in which the region containing the black hole is described by a Vaidya solution) but this too is artificial because all the photons have to become radially directed within some point and it also violates the 2nd law of thermodynamics.

Apart from these examples, it appears that there is no self-similar solution containing a black hole in either an exact or asymptotically Friedmann background for any value of  $\gamma$  in the range  $1 \leq \gamma \leq 2$ . There are only asymptotically quasi-Friedmann solutions, including “universal” black holes without a black hole event horizon or cosmological particle horizon. However, there are self-similar asymptotically Friedmann solutions for  $0 < \gamma < 2/3$ , which suggests that PBHs *can* grow as fast as the universe in the presence of dark energy (at least for a limited period). This conclusion may also apply for a quintessence field, although this has not been rigorously proved.

The difference between the positive and negative pressure solutions is important in two respects. First, while the negative-pressure ones are physically well-motivated in the inflationary scenario, because one might expect the associated density perturbations to extend to “infinity”, the positive-pressure ones are theoretically unmotivated and may also be observationally excluded for some parameter because they have a solid angle deficit at large distances. Second, self-similar black holes only exist if their size as a fraction of the cosmological horizon is not too small in the positive-pressure case but not too large in the negative-pressure situation. This means that less fine-tuning is required in the latter case. It is interesting that there is no accretion in the limit  $\gamma \rightarrow 0$ , which is consistent with the Schwarzschild-de Sitter solution.

The existence of these self-similar black hole solutions suggests that black holes may increase their mass by a considerable factor during any dark-energy or quintessence dominated era, regardless of whether or not they are “primordial” [16]. There are two contexts in which this effect may be important: (1) in the period immediately after any PBHs formed at the end of inflation; (2) in the recent period when whatever dominates the density of the universe causes it to accelerate. Bean and Magueijo [16] focussed on the first situation but the second one may also be interesting. For although the accretion factor may not be very large up to now, the black hole mass will continue to grow like cosmic time so as long as the dark energy dominates the cosmological density. In the simplest models, this applies indefinitely, so the black hole can grow arbitrarily large. We are investigating the astrophysical implications of this result – and its implications for the Bean-Magueijo claim that PBHs can grow large enough to provide the supermassive black holes in galactic nuclei – in a separate paper.

Finally, we note that our analysis restricts the situations in which the similarity hypothesis [81] applies. This hypothesis claims that there are certain circumstances in which spherically symmetric solutions evolve to self-similar form. The present work shows that there are at least some situations in which this does *not* happen.

(Even if there were a self-similar solution, one would still need to show that it was stable in order for it to be an attractor [81].) On the other hand, the similarity hypothesis *does* appear to hold in spherical gravitational collapse when the pressure is positive but very small ( $0 < \gamma - 1 \ll 1$ ) [82, 83]. It may also hold in the negative-pressure situation.

### Acknowledgments

The authors would like to thank R.H. Henriksen, H. Motohashi and M. Nakashima for useful input. TH was supported by the Grant-in-Aid for Scientific Research Fund of the Ministry of Education, Culture, Sports, Science and Technology, Japan (Young Scientists (B) 18740144 and 21740190). TH was also grateful to CECS for its hospitality during his visit by the Fondecyt grant 7080214. HM was supported by Fondecyt grant 1071125. The Centro de Estudios Científicos (CECS) is funded by the Chilean Government through the Millennium Science Initiative, the Centers of Excellence Base Financing Program of Conicyt, and Conicyt grant “Southern Theoretical Physics Laboratory” ACT-91. CECS is also supported by a group of private companies which at present includes Antofagasta Minerals, Arauco, Empresas CMPC, Indura, Naviera Ultragas, and Telefónica del Sur. BJC thanks the Research Center for the Early Universe at the University of Tokyo for hospitality received during this work. This work has benefited from exchange visits supported by a JSPS and Royal Society bilateral grant.

### Appendix A: Physical dimensions

We take the action to be

$$S = \int dt d^3x \sqrt{-g} \left[ \frac{c^4}{16\pi G} \mathcal{R} - \left( \frac{1}{2} \phi_{,\mu} \phi^{,\mu} + V(\phi) \right) \right], \quad (\text{A1})$$

where the coordinates in 4-dimensional Lorentzian space-time are  $(x^0, x^i) = (ct, x^i)$  with  $i = 1, 2, 3$ . We emphasize that  $dt d^3x$  must not be written as  $d^4x$  if we include factors of  $c$  explicitly. The Einstein equations are

$$G^\mu{}_\nu = \frac{8\pi G}{c^4} T^\mu{}_\nu, \quad (\text{A2})$$

$$T_{\mu\nu} = \left( \phi_{,\mu} \phi_{,\nu} - \frac{1}{2} g_{\mu\nu} \phi_{,\rho} \phi^{,\rho} \right) - g_{\mu\nu} V. \quad (\text{A3})$$

The dimensions of the action and gravitational constant are

$$[S] = ML^2 T^{-1}, \quad (\text{A4})$$

$$[G] = M^{-1} L^3 T^{-2}, \quad (\text{A5})$$

so we obtain

$$[c^4/G] = MLT^{-2}, \quad (\text{A6})$$

$$[dt d^3x \times c^4/G] = ML^4 T^{-1}, \quad (\text{A7})$$

and hence

$$[T_{\mu\nu}] = ML^{-1} T^{-2}. \quad (\text{A8})$$

The relativistic factor in the black hole metric is

$$[GM/c^2] = L. \quad (\text{A9})$$

For a perfect fluid, the energy momentum tensor is

$$T_{\mu\nu} = (\mu + p) u^\mu u^\nu + p g_{\mu\nu}, \quad (\text{A10})$$

where  $u^\mu \equiv dx^\mu/d(c\tau)$  is the dimensionless 4-velocity of the fluid element, with  $\tau$  being the affine time. The mass density  $\rho$ , energy density  $\mu$  and pressure  $p$  have dimensions

$$[\rho] = ML^{-3}, \quad (\text{A11})$$

$$[\mu] = ML^{-1} T^{-2}, \quad (\text{A12})$$

$$[p] = ML^{-1} T^{-2}, \quad (\text{A13})$$

so

$$[G\mu/c^2] = [Gp/c^2] = [G\rho] = T^{-2}. \quad (\text{A14})$$

For a scalar field, we have

$$[\phi] = [c^2/\sqrt{G}] = M^{1/2} L^{1/2} T^{-1}, \quad (\text{A15})$$

$$[V] = [c^4/(GL^2)] = ML^{-1} T^{-2}. \quad (\text{A16})$$

Hence the exponential potential can be written in the dimensionally correct form

$$V(\phi) = V_0 e^{-\sqrt{8\pi G} \lambda \phi / c^2}, \quad (\text{A17})$$

where  $\lambda$  is a dimensionless constant. The equation of motion for  $\phi$  is

$$\square\phi = \frac{\partial V}{\partial\phi}. \quad (\text{A18})$$

In the cosmological context, a scalar field is equivalent to a perfect fluid with

$$\mu = \frac{1}{2} \left( \frac{d\phi}{dx^0} \right)^2 + V, \quad (\text{A19})$$

$$p = \frac{1}{2} \left( \frac{d\phi}{dx^0} \right)^2 - V. \quad (\text{A20})$$

For the flat Friedmann metric

$$ds^2 = -c^2 dt^2 + (t/t_1)^{4/\lambda^2} dl_3^2, \quad (\text{A21})$$

where  $dl_3$  is the flat line element, and

$$T^t{}_t = -\mu = -\frac{3c^2}{2\pi G \lambda^4 t^2}. \quad (\text{A22})$$

This is the energy density for a Friedmann universe filled by a scalar field with an exponential potential. The last two equations also apply for a perfect fluid with  $p = (\gamma - 1)\mu$  if one replaces  $\lambda$  with  $\sqrt{3\gamma}$ .



## Appendix B: Energy conditions

We summarize the energy conditions for a matter field with energy-momentum tensor given in the diagonal form as  $T^\mu{}_\nu = \text{diag}(-\mu, p_r, p_t, p_t)$  [84, 85]. The physical interpretations of  $\mu$ ,  $p_r$  and  $p_t$  are the energy density, radial pressure and tangential pressure, respectively.

The inequality  $T_{\mu\nu}W^\mu W^\nu \geq 0$  for any timelike vector  $W^\mu$  is called the weak energy condition (WEC). This guarantees that a timelike observer measures a non-negative energy density and implies  $\mu \geq 0$ ,  $p_r + \mu \geq 0$  and  $p_t + \mu \geq 0$ . The null energy condition (NEC) just replaces  $W^\mu$  with a future-directed null vector in the above definition, which means that one can drop the condition  $\mu \geq 0$ .

The condition that  $T^{\mu\nu}W_\nu$  is a future-directed non-spacelike vector for every future-directed non-spacelike vector  $W^\mu$  is called the dominant energy condition

(DEC). This means that the mass-energy can never flow faster than the speed of light and implies  $\mu \geq 0$ ,  $-\mu \leq p_r \leq \mu$  and  $-\mu \leq p_t \leq \mu$ .

The inequality  $(T_{\mu\nu} - g_{\mu\nu}T^\rho{}_\rho/2)W^\mu W^\nu \geq 0$  for any timelike vector  $W^\mu$  is called the strong energy condition (SEC). If there is no cosmological constant, this assures the timelike convergence condition  $R_{\mu\nu}W^\mu W^\nu \geq 0$  in general relativity, which means that the gravitational force is essentially attractive. The SEC implies  $p_r + \mu \geq 0$ ,  $p_t + \mu \geq 0$  and  $p_r + 2p_t + \mu \geq 0$ .

For  $p_r = p_t \equiv p$ , as assumed throughout this paper, the WEC implies  $\mu \geq 0$  and  $p + \mu \geq 0$ , the DEC implies  $\mu \geq 0$  and  $-\mu \leq p \leq \mu$ , the SEC implies  $p + \mu \geq 0$  and  $3p + \mu \geq 0$ , and the NEC implies  $\mu + p \geq 0$ . The third and fourth conditions must be violated for inflation and the dynamical wormholes, respectively. Note that DEC implies WEC, WEC implies NEC, and SEC implies NEC, but SEC does not imply WEC.

- 
- [1] Ya.B. Zel'dovich and I.D. Novikov, *Sov. Astron. A. J.* **10**,602, (1967).
  - [2] S.W. Hawking, *Mon. Not. R. Astron. Soc.* **152**,75 (1971).
  - [3] B.J. Carr and S.W. Hawking, *Mon. Not. R. Astron. Soc.* **168**, 399 (1974).
  - [4] S.W. Hawking, *Commun. Math. Phys.* **43**, 199 (1975).
  - [5] B.J. Carr, Ph.D. thesis, Cambridge University (1976).
  - [6] G.V. Bicknell and R.N. Henriksen, *Astrophys. J.* **225**, 237 (1978).
  - [7] B.J. Carr and A. Yahil *Ap.J.* **360**, 330 (1990).
  - [8] H. Maeda, J. Koga and K-i. Maeda, *Phys. Rev.* **D66**, 087501 (2002).
  - [9] D. Lin, B.J. Carr, and S.M. Fall, *Mon. Not. R. Astron. Soc.* **177**, 51 (1976).
  - [10] G.V. Bicknell and R.N. Henriksen, *Astrophys. J.* **219**, 1043 (1978).
  - [11] M. S. Madsen *Class. Quant. Grav.* **5**, 627 (1988).
  - [12] T. Harada and B.J. Carr, *Phys. Rev.* **D71**, 104010 (2005).
  - [13] T. Harada, H. Maeda and B.J. Carr, *Phys. Rev.* **D74**, 024024 (2006).
  - [14] S. Perlmutter *et al.*, *Astrophys. J.* **517**, 565 (1999); A.G. Riess *et al.*, *Astron. J.* **116**, 1009 (1998); *Astron. J.* **117**, 707 (1999).
  - [15] J. Kormendy and D. Richstone, *Ann. Rev. Astron. Astrophys.* **33**, 58 (1995).
  - [16] R. Bean and J. Magueijo, *Phys. Rev.* **D66**, 063505 (2002).
  - [17] T. Harada, H. Maeda and B.J. Carr, *Phys. Rev.* **D77**, 024022 (2008).
  - [18] H. Maeda, T. Harada, B.J. Carr, *Phys. Rev.* **D77**, 024023 (2008).
  - [19] A. Einstein and E.G. Strauss, *Rev. Mod. Phys.* **17**, 120 (1945).
  - [20] K.R. Nayak, M.A.H. MacCallum and C.V. Vishveshwara, *Phys. Rev.* **D63**, 024020 (2000).
  - [21] J. Sultana and C.C. Dyer, *Gen. Rel. Grav.* **37**, 1349 (2005).
  - [22] D. Kastor and J.H. Traschen, *Phys. Rev.* **D47**, 5370 (1993).
  - [23] K. Maeda, N. Ohta and K. Uzawa, *JHEP* **0906**, 052 (2009).
  - [24] G.W. Gibbons and K. Maeda, arXiv:0912.2809 [gr-qc] (2009).
  - [25] B. J. Carr and A. A. Coley, *Class Quant. Grav.* **16** R31-R71 (1999).
  - [26] H. Bondi, *Mon. Not. R. Astron. Soc.* **112**, 195 (1952).
  - [27] B.J. Carr, *Ap. J.* **201**, 1-19 (1975).
  - [28] B.J. Carr and M. Sakellariadou, *Ap. J.* **516**, 195 (1999).
  - [29] T. Harada and B.J. Carr, *Phys. Rev.* **D71**, 104009 (2005).
  - [30] P. Coles, *Cosmology: The Origin and Evolution of Cosmic Structures* (John Wiley, New York, 1995), p. 212.
  - [31] E. Babichev, V. Dokuchaev, Yu. Eroshenko, *J. Exp. Theor. Phys.* **100**, 528 (2005).
  - [32] M.E. Cahill and A.H. Taub, *Commun. Math. Phys.* **21**, 1 (1971).
  - [33] G.V. Bicknell and R.N. Henriksen, *Astrophys. J.* **232**, 670 (1979).
  - [34] D. K. Nadezhin, I. D. Novikov and A. G. Polnarev, *Sov. Astron.* **22**, 129 (1978).
  - [35] I. D. Novikov and A. G. Polnarev, *Sov. Astron.* **24**, 147 (1980).
  - [36] B.J. Carr and C. Gundlach, *Phys. Rev.* **D67**, 024035 (2003).
  - [37] A. Ori and T. Piran, *Phys. Rev.* **D42**, 1068 (1990).
  - [38] T. Foglizzo and R.N. Henriksen, *Phys. Rev.* **D48**, 4645 (1993).
  - [39] M. Barriola and A. Vilenkin, *Phys. Rev. Lett.* **63**, 341 (1989).
  - [40] V. Perlick, *Living Rev. Rel.* **7**, 9 (2004).
  - [41] B.J. Carr and A. Coley, *Phys. Rev.* **D62**, 044023 (2000).
  - [42] M. Goliath, U. Nilson and C. Ugglä, *Class. Quant. Grav.* **15**, 167 (1998); **15**, 2841 (2000).
  - [43] J. C. Niemeyer and K. Jedamzik, *Phys. Rev.* **D59**, 124013 (1999).
  - [44] K. Jedamzik and J. C. Niemeyer, *Phys. Rev.* **D59**, 124014 (1999).
  - [45] M. Shibata and M. Sasaki, *Phys. Rev.* **D60**, 084002 (1999).
  - [46] I. Hawke and J. M. Stewart, *Class. Quant. Grav.* **19**, 3687 (2002).
  - [47] I. Musco, J. C. Miller and L. Rezzola, *Class. Quant. Grav.* **22**, 1405 (2005).
  - [48] B.J. Carr, A. Coley, M. Goliath, U. Nilson and C. Ugglä, *Phys. Rev.* **D61**, 081502 (2000); *Class. Quant. Grav.* **18**, 303-324 (2001).

- [49] B.C. Reed and R.N. Henriksen, *Astrophys. J.* **236**, 338 (1980).
- [50] S. Hacyan, *Ap.J.* **229**, 42 (1979).
- [51] P. R. Brady, *Phys. Rev. D***51**, 4168 (1995).
- [52] S. Bhattacharya, P.S. Joshi and K. Nakao, arXiv: 0911.2297 (2009).
- [53] M.D. Roberts, *Gen. Rel. Grav.* **21**, 907 (1989).
- [54] Y. Oshiro, K. Nakamura and A. Tomimatsu, *Prog. Theor. Phys.* **91**, 1265 (1994).
- [55] I.I. Gutman and R.M. Bepal'ko, *Sbornik Sovrem. Probl. Grav. Tbilissi*, **1**, 201 (1967).
- [56] H. Maeda, *Phys. Rev. D* **79**, 024030 (2009).
- [57] B. Ratra and J. Peebles, *Phys. Rev. D***52**, 1837 (1995).
- [58] P. S. Custodio and J. E. Horvath, *Int. J. Mod. Phys. D***14**, 257 (2005).
- [59] T. Jacobson, *Phys. Rev. Lett.* **83**, 2699 (1999). **93**, 021102 (2004).
- [60] A. Frolov and L. Kofman, *JCAP* **5**, 9 (1999).
- [61] R. Akhoury, C. Gauthier and A. Vikman, *JHEP* **03**, 082 (2009).
- [62] M. Kyo, T. Harada and H. Maeda, *Phys. Rev. D***77**, 124036 (2008).
- [63] E. Babichev, V. Dokuchaev, Yu. Eroshenko, *Phys. Rev. Lett.* **93**, 021102 (2004).
- [64] B. J. Carr and A. A. Coley, *Class Quantum Grav.* **17** 4339 (2000).
- [65] B.J. Carr and A. Koutras *Ap.J.* **405**, 34 (1993).
- [66] A. Nusser, *Mon. Not. R. Astron. Soc.* **375**, 1106 (2007).
- [67] H. Maeda, T. Harada, B.J. Carr, *Phys. Rev. D***79**, 044034 (2009).
- [68] M.S. Morris and K.S. Thorne, *Am. J. Phys.* **56**, 395 (1988).
- [69] M. Visser, *Lorentzian Wormholes: From Einstein to Hawking*, (Springer-Verlag, Berlin, Germany, 1997).
- [70] D. Hochberg and M. Visser, *Phys. Rev. D***56**, 4745 (1997).
- [71] D. Hochberg and M. Visser, *Phys. Rev. D***58**, 044021 (1998).
- [72] S.A. Hayward, *Int. J. Mod. Phys. D***8**, 373 (1999).
- [73] S.A. Hayward, *Phys. Rev. D***49**, 6467 (1994).
- [74] E. Babichev, V. Dokuchaev, Yu. Eroshenko, *Class. Quant. Grav.* **22**, 143 (2005).
- [75] R.R. Caldwell, *Phys. Lett. B***545**, 23 (2002).
- [76] R.R. Caldwell, M. Kamionkowski and N.N. Weinberg, *Phys. Rev. Lett.* **91**, 071301 (2003).
- [77] A. Frolov, *Phys. Rev. D***70**, 061501 (2004).
- [78] S. Mukohyama, *Phys. Rev. D***71**, 104019 (2005).
- [79] E. Babichev, V. Mukhanov and A. Vikman, *JHEP* **02**, 101 (2008).
- [80] E. Babichev, V. Mukhanov and A. Vikman, *JHEP* **09**, 061 (2006).
- [81] B.J. Carr & A.A. Coley, *Gen. Rel. Grav.* **17**, 4339 (2000).
- [82] T. Harada and H. Maeda, *Phys. Rev. D***63**, 084022 (2001).
- [83] M. Snajdr, *Class. Quant. Grav.* **23**, 3333 (2006).
- [84] S.W. Hawking and G.F.R. Ellis, *The Large scale structure of space-time* (Cambridge University Press, Cambridge, 1973).
- [85] S.M. Carroll, *Spacetime and Geometry: An Introduction to General Relativity* (Addison-Wesley, San Francisco, 2004).