



Relative velocities among accreting planetesimals in binary systems: the circumbinary case

H. Scholl, F. Marzari, P. Thébault

► To cite this version:

H. Scholl, F. Marzari, P. Thébault. Relative velocities among accreting planetesimals in binary systems: the circumbinary case. *Monthly Notices of the Royal Astronomical Society*, 2007, 380, pp.1119-1126. <10.1111/j.1365-2966.2007.12145.x>. <hal-00388161>

HAL Id: hal-00388161

<https://hal.science/hal-00388161v1>

Submitted on 14 Dec 2020

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.



HAL Authorization

Relative velocities among accreting planetesimals in binary systems: the circumbinary case

H. Scholl,¹★ F. Marzari² and P. Thébault^{3,4}

¹*Laboratoire Cassiopée, CNRS, Observatoire de la Côte d'Azur, B.P. 4229, F-06304 Nice, France*

²*Department of Physics, University of Padova, Via Marzolo 8, 35131 Padova, Italy*

³*Stockholm Observatory, Albanova Universitetcentrum, SE-10691 Stockholm, Sweden*

⁴*Observatoire de Paris, Section de Meudon, F-92195 Meudon Principal Cedex, France*

Accepted 2007 June 22. Received 2007 June 15; in original form 2007 May 16

ABSTRACT

We numerically investigate the possibility of planetesimal accretion in circumbinary discs, under the coupled influence of both stars' secular perturbations and friction due to the gaseous component of the protoplanetary disc. We focus on one crucial parameter: the distribution of encounter velocities between planetesimals in the 0.5–100 km size range. An extended range of binary systems with differing orbital parameters is explored. The resulting encounter velocities are compared to the threshold velocities below which the net outcome of a collision is accumulation into a larger body instead of mass erosion. For each binary configuration, we derive the critical radial distance from the binary barycentre beyond which planetesimal accretion is possible. This critical radial distance is smallest for equal-mass binaries on almost circular orbits. It shifts to larger values for increasing eccentricities and decreasing mass ratio. The importance of the planetesimals' orbital alignments of planetesimals due to gas drag effects is discussed.

Key words: accretion, accretion discs.

1 INTRODUCTION

At present more than 40 planets have been found in binary star systems, which corresponds to about 20 per cent of the whole sample of known extrasolar planets (Desidera & Barbieri 2007). All of them are on so-called S-type or circumprimary orbits that encircle one component of the double star system. Planets may also revolve on dynamically stable orbits about both binary stars on so-called P-type or circumbinary orbits (e.g. Holman & Wiegert 1999). At present, there is only one planet assumed to revolve on a circumbinary orbit. The primary, HD 202206, is a metal-rich main-sequence star. Its companion, HD 202206b, is assumed to be a low-mass brown dwarf with an estimated mass of $17.5 M_{\text{Jup}}$ according to radial velocity measurements (Udry et al. 2002; Correia et al. 2005). This mass is beyond the widely accepted brown dwarf limit of $10 M_{\text{Jup}}$. The third body, HD 202206c, is assumed to be a Jupiter-like planet revolving about the binary. Adopting a mass of $2.41 M_{\text{Jup}}$ for HD 202206c, Correia et al. (2005) show in a dynamical analysis that HD 202206c and HD 202206b might be in a 5/1 mean motion resonance. This resonant configuration is possibly a consequence of an inward migration of HD 202206c due to the forces exerted by a viscous circumbinary disc (Nelson 2003).

The signature of circumbinary planets in radial velocity measurements is in most cases difficult to be spotted because of the short-term large-amplitude velocity fluctuations on the primary induced by the companion star. However, planets might be very frequent among the population of close star couples. Mid-infrared emissions have in fact revealed substantial circumbinary material around pre-main-sequence close binaries like DQ Tau, UZ Tau, GW Ori (Mathieu et al. 2000) and AK Sco (Jensen & Mathieu 1997), and more detached binaries like GG Tau (Dutrey, Guilloteau & Simon 1994), UY Aur (Close et al. 1998). These discs are significantly more massive than the minimum-mass solar nebula which suggests abundant planetary formation. However, both concurrent mechanisms for extrasolar planet formation, namely the standard solid-core model forming planetary embryos from runaway and oligarchic accretion of planetesimals (e.g. Greenberg et al. 1978; Wetherill & Stewart 1989; Kokubo & Ida 2000) possibly followed by gas infall on to the core for the formation of giant planets (Pollack et al. 1996; Bodenheimer & Lin 2002), and the alternative model of local gravitational collapse of disc material (Boss 1997) might be affected by the secular perturbations of the binary.

We will focus here specifically on the 'standard' solid-core scenario, in particular, on the stage of planetary embryo accretion. This stage is just before the one investigated by Quintana & Lissauer (2006) who determined regions around close binaries where planets can form by accumulation of embryos. Within the frame of this

★E-mail: hans.scholl@oca.eu

model, the initial stage of planetesimal accretion is the runaway growth of isolated planetary embryos. This stage is very fast and efficient, provided that the encounter velocities among planetesimal are low, that is, much smaller than the escape velocities of the growing embryos (Lissauer 1993). As a consequence, runaway growth is particularly sensitive to any external perturbation able to stir up relative velocities in the planetesimal swarm. In the specific case of a circumbinary disc, relative velocities between planetesimals can be expected to be higher as compared to the single-star case. In the vicinity of the binary, velocities could be increased to values for which runaway accretion is no longer possible, or even to higher values for which collisions may result in mass removal via cratering or fragmentation rather than accretion. The threshold velocity for the stop of runaway is of the order of the escape velocities $v_{\text{esc}(R)}$ of the growing bodies while the threshold velocity v_{ero} for the limit between accreting and eroding impacts is less straightforward to estimate and depends on the respective sizes and physical properties of the two colliding planetesimals. In any case, the distribution of encounter velocities is the key parameter controlling the fate of a collisionally interacting swarm of planetesimals. This is the issue which we address here, by numerically estimating the distribution of $\langle \Delta v \rangle$ in a disc of planetesimals surrounding a binary system in order to derive assertive conclusions on the possibility of planet formation.

Only taking into account the gravitational potential of both stars might lead to incorrect estimates of relative velocities, in particular close to the star couple. Indeed, the planetesimal accretion phase is believed to take place while a large fraction of the primordial gas disc is still present. Gas drag is known to circularize orbits (Adachi, Hayashi & Nakazawa 1976) and thus reduce relative velocities in dynamically hot systems (Marzari & Scholl 2000). Besides circularization, gas drag has the additional effect of aligning orbital apsides, an effect which is also called orbital synchronization (Marzari et al. 1997; Marzari & Scholl 2000). The consequences of this synchronization on relative velocities depend on planetesimal sizes. Bodies of the same size all have their periaapses aligned in the same direction. Smaller bodies have a much smaller dispersion of periaapse directions than larger bodies. Furthermore, the mean direction of the periastron alignment is also different for objects of different sizes (Thébault, Marzari & Scholl 2006). As a result, while relative velocities between small planetesimals of the same size may drop to almost zero, relative velocities among bigger bodies are more significant because of a larger periaapse direction dispersion. The highest relative velocities can be expected between objects of *different* sizes due to the size dependency of the periastron alignment. The relative velocity distribution depends, therefore, not only on the radial distance to the barycentre of the binary but also on the sizes of colliding planetesimals and on the gas density.

In a pioneering attempt to model planetary accretion in circumbinary discs, Moriwaki & Nakagawa (2004) used a purely gravitational model without gas drag. They found very stringent limits on the minimum distance from the binary beyond which planetesimal accumulation is possible. However, as pointed out above, neglecting the velocity damping effect and orbit alignment due to gas friction may yield misleading results for the efficiency of planetesimal accumulation, in particular in the vicinity of the binary. In the present work, we redress this issue by performing numerical simulations where the effects of gas drag are taken into account, although with unavoidable simplifying assumptions. Moreover, we compute the average encounter velocities between planetesimals from their trajectories and not from the values of forced eccentricities as Moriwaki & Nakagawa (2004) who used the relation $\Delta v \simeq e v_{\text{kep}}$, where v_{kep}

denotes Keplerian velocity. This often used relation holds only under the condition that all orbits have fully randomized Keplerian angles. However, in circumbinary discs with gas, orbital phasing of periastra directions is of high importance because of the combined effects of secular perturbations and gas friction. Another limitation of the $\Delta v \simeq e v_{\text{kep}}$ relation is that it works *locally* and it does not take into account the possibility of large radial excursions of bodies with high eccentricities (for more on this subject, see e.g. the discussion in Thébault & Doressoundiram 2003). Both these issues suggest that the relation used by Moriwaki & Nakagawa (2004) may not be appropriate in presence of large orbital eccentricities and possibly of gas drag. Our numerical approach, which takes into account all close particle encounters representing collisions, has the advantage to handle automatically orbital phasing (and dephasing) as well as radial mixing effects. This allows us to derive more reliable values for the limiting distance from the barycentre of the binary where planet formation is possible for different parameters of the binary. Not surprisingly, our results significantly differ from those of Moriwaki & Nakagawa (2004) in particular close to the binary where planetary accretion is more likely to form planets due to the higher density of material.

Due to obvious numerical constraints, mutual gravitational interactions among planetesimals, like for instance dynamical friction, are not included even if, in principle, they might be taken into account in N -body simulations. However, the speed of current computers is still far too low to allow a systematic investigation of binary systems surrounded by a self-gravitating planetesimal disc. These simulations are still restricted to discs with massless planetesimals. Their size distribution, of course, does not evolve with time. To do so, particle-in-a-box computations, which can follow the evolution of a large number of massive bodies, are required, but the price to pay is a significant loss of accuracy in the computations of the orbital evolution of the planetesimals that, in the present scenario, is crucial. As a consequence, we decided to adopt a deterministic numerical approach, which, although neglecting important physical effects, gives a reliable estimate of encounter velocity distributions and thus enable to derive ‘maps’ of the disc regions with relative velocities low enough to favour planetesimal accretion and planet formation.

Our numerical model is described in detail in Section 2. Section 3 is devoted to the computations performed in the gas-free case, a scenario similar to that of Moriwaki & Nakagawa (2004). In Section 4 we use the results obtained from models including gas drag in order to map the regions where we expect planet formation by accretion. Section 5 is dedicated to the discussion of our results.

2 THE NUMERICAL MODEL

2.1 Main characteristics

Our deterministic code is an upgraded version from the one used in several earlier studies (Marzari & Scholl 1998, 2000), adapted to the specificities of the circumbinary configuration. We consider a system of $N \sim 10\,000$ planetesimals, treated as non-gravitationally interacting test particles, submitted to the gravitational potentials of both stars and to a gas drag force following Adachi et al. (1976) and described in the following section. Planetesimal orbits are integrated using a fourth-order Runge–Kutta scheme with fixed step size. We like to point out that for the circumpriary problem (Thébault et al. 2006) we applied a different numerical scheme which is less convenient for the circumbinary case.

The binary stars move on an elliptic orbit of eccentricity e_b and semimajor axis a_b . For sake of simplicity and to limit the number of free parameters to a manageable value, we assume for all runs $a_b = 1$ au. This is the standard value also taken by Moriwaki & Nakagawa (2004) in their numerical explorations. We are thus here implicitly focusing on planetesimal accretion in the regions typically beyond 3–4 au from the binary’s centre of mass. This region is sufficiently away from the critical instability region found by Holman & Wiegert (1999) to guarantee orbital stability for almost planar and circular motion. Both stars’ total mass is fixed and equal to $1 M_\odot$. The primary parameters of our model are thus the mass ratio and eccentricity of the binary star system. We use five different mass ratios q for the binary, $q = 0.5, 0.4, 0.3, 0.2$ and 0.1 . The stellar masses m_1 of the primary and m_2 of the companion are then $m_1 = 1 - q$ and $m_2 = q$, in M_\odot units. For each mass ratio, models with six different binary eccentricities e_b are simulated with $e_b = 0.0, 0.1, 0.2, 0.3, 0.4$ and 0.5 . This gives a total of 30 models. Each model is simulated with and without gas drag. In models taking into account gas drag, we consider planetesimal of physical radii ranging from 0.5 to 100 km. The initial eccentricities and orbital inclinations are chosen with a uniform random distribution between 0 and 10^{-5} . The initial particle semimajor axes are randomly distributed between $a_{\min} = 4$ au and $a_{\max} = 12$ au. The value for a_{\min} being approximately equal to the inner truncation of a circumbinary disc by the tidal influence of the binary (Artymowicz & Lubow 1994; Gunther & Kley 2002). Whenever the semimajor axis of a planetesimal decreases below the inner limit because of drag friction, the body is removed.

The typical time-scale for one run is $\sim 100\,000$ binary revolutions, which is approximately the time necessary to reach a steady state, with fully developed secular perturbations and the gas drag induced orbital phasing. After such a steady state is reached, we determine relative velocities among the planetesimal population. For this purpose, we use the positions and velocities of all planetesimals at a fixed time. For each planetesimal, we search its nearest neighbouring planetesimal and compute the relative velocity corrected for Kepler shear due the different radial distances. Relative velocities are sampled in bins in radial direction. A bin size of 0.5 au is used. The median value for each bin is then taken as the representative value for the collisional or impact velocity of planetesimals at the corresponding radial distance. Since, as pointed out, the orbit alignment is size dependent, we compute the median relative velocities $\Delta v_{(R1,R2)}$ between planetesimal populations for each pairs of radii R_1, R_2 in each bin. We typically use 20–30 close encounters in a bin to determine the corresponding median relative velocity.

2.2 The gas drag force

The drag force is modelled in laminar gas approximation by

$$\dot{\mathbf{r}} = -K v_{\text{rel}} \mathbf{v}_{\text{rel}}, \quad (1)$$

where K is the drag parameter given by (Kary, Lissauer & Greenzweig 1993) as

$$K = 3 \frac{\rho_g C_d}{8 \rho_{\text{pl}} s}, \quad (2)$$

where s is the planetesimal radius, ρ_{pl} its mass density, ρ_g the gas density of the protoplanetary disc and C_d a dimensionless drag coefficient related to the shape of the body (~ 0.4 for spherical bodies). Exploring the gas drag density profile as a free parameter would be too CPU time consuming and we shall restrict ourselves to one ρ_g

distribution. We assume the standard minimum-mass solar nebula (MMSN) of Hayashi (1981), with $\rho_g = \rho_{g0}(a/1 \text{ au})^{-2.75}$ and $\rho_{g0} = 1.4 \times 10^{-9} \text{ g cm}^{-3}$. We take a typical value $\rho_{\text{pl}} = 3 \text{ g cm}^{-3}$.

We are here implicitly assuming that the gas disc is axisymmetric and pressure supported. We are aware that this axisymmetric prescription does probably not represent the real physical structure of the circumbinary gaseous component of the disc. This is a complex issue, which has only been addressed in a handful of elaborated numerical investigations (e.g. Artymowicz & Lubow 1994; Gunther & Kley 2002; Nelson 2003; Pichardo, Sparke & Aguilar 2005). All these studies clearly show the inner tidal truncation of the disc, but are less detailed on the structure within the disc, which is particularly complex to be precisely computed, especially in the eccentric binary case. The only clear result is the onset of complex time-evolving spiral patterns extending from the inner edge of the disc towards the outer regions. At present it appears a very demanding task to perform a fully coupled numerical treatment of both the gas disc and an embedded planetesimal swarm. For this reason, we prefer in a first step a simplified approach where the gas drag force is just described by equation (1). The onset of density waves in the disc induces radial and azimuthal variations both of the disc density and of the gas velocity where an additional radial component adds up to the Keplerian tangential component. These effects might alter the strength of the drag acceleration when a planetesimal enters and exits the higher or lower density region. However, it is reasonable to assume that, on the average, this effect would not strongly modify accelerations obtained by equation (1) and thus the dynamical evolution of planetesimals, at least for planetesimals in the km-size range. In particular, we expect the occurrence of the periastron alignment.

As already mentioned, we start all planetesimals on almost circular orbits. This choice is from a dynamical point of view equivalent to assuming that, at $t = 0$, the proper eccentricity is almost equal to the forced one, e_f . This is a somewhat artificial and not a necessarily realistic starting configuration motivated by the fact that, at present, we have no generally accepted clues either on how planetesimals form within a disc with spiral waves or on the initial orbital elements of the planetesimals as they detach from the gas. We integrate the model until a steady state is reached where periapses are aligned and the interaction with the gas becomes only a minor perturbation of the Keplerian orbit. This procedure is valid for small planetesimals, since they reach the same periastron alignment independently of their starting values with the same distribution of proper eccentricities. Larger planetesimals, which are much less affected by gas drag, have their proper eccentricities increasing during their growth by collisions and close encounters with other large planetesimals. As a consequence, they too should lose memory of their initial orbits. As a consequence, we think that our initial choice of eccentricities does not introduce too significant effects on the calculation of relative impact velocities once the steady state is reached.

3 IMPACT VELOCITIES IN THE GAS-FREE CASE

In the gas-free models we use 25 000 particles with semimajor axes ranging from 3 to 50 au. However, we focus here essentially on the region within 12 au, since, as it clearly appears in Fig. 1, average encounter velocities drop down to almost zero beyond this limit. For the $r < 12$ au region, the $\langle \Delta v \rangle$ we obtain might be fitted, for the range of stellar masses are orbital eccentricities explored, by the

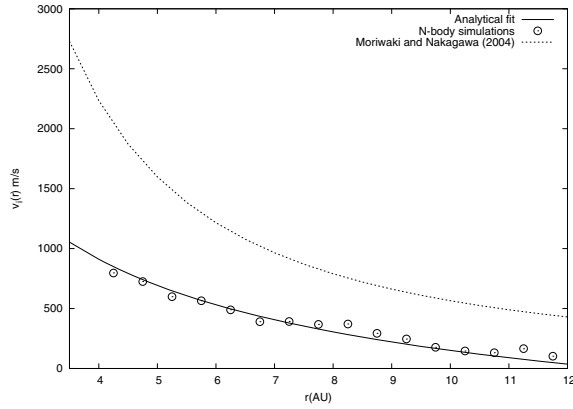


Figure 1. Relative impact velocities between planetesimals computed with our equation (3) (continuous line) compared to the predictions of Moriwaki & Nakagawa (2004) (dashed line). The empty circles show the numerical results of the N -body integrations. The parameters of the binary are $q = 0.2$ and $e_b = 0.4$.

empirical formula

$$\begin{aligned} \langle \Delta v \rangle_{(r)} = & (-384.2 + 2067.7q - 447522e_b \\ & - 3864.9q^2 + 9462.4qe_b) \\ & + (1024.8 - 5783.4 + 16431e_b \\ & + 11251q^2 - 34396qe_b) \frac{1}{\sqrt{r}} \end{aligned} \quad (3)$$

which is a second-order polynomial in q and e_b multiplied by an additional term equal to \sqrt{r} accounting for the radial dependence of $\langle \Delta v \rangle$ on the local Keplerian velocity.

Since we neglect the gas drag force, the radius of the planetesimals is not a parameter of the problem. In Fig. 1, we also compare our results to the $\langle \Delta v \rangle$ computed with the Moriwaki & Nakagawa (2004) analytical approximation for the specific case with $q = 0.2$ and $e_b = 0.4$. For reference, we also plot the numerical data obtained from the N -body integration. It can be clearly seen that our numerically derived values are always significantly smaller than the Moriwaki & Nakagawa (2004) estimates. This shows the weakness of the approximate relation $\Delta v \simeq ev_{\text{kep}}$ for the present case, where orbital phasing and radial excursions of the planetesimals play a crucial role in determining the average impact velocity. As an example, while our simulations give low relative velocities, of the order of 50 m s^{-1} at about 11 au from the barycentre, the analytical approximation of Moriwaki & Nakagawa (2004) predicts such low velocities only beyond 50 au. Our numerical results give much lower values for impact velocities which allow planetary formation to occur much closer to the barycentre of the binary even in absence of gas friction. This is a first indication that planets on circumbinary orbits may be abundant due to the low relative velocities in the regions near the binary with higher surface density. In the following section, we show how this promising result is affected by gas drag.

4 IMPACT VELOCITIES IN THE PRESENCE OF GAS

We now turn to the more realistic case including gas drag. As previously mentioned, it is well known that the combination of the drag force and perturbations by the gravity field of the binary causes a strong periairapse phasing of the orbits of the planetesimals. The planetesimals' physical sizes are now a crucial parameter. Bodies of

similar sizes have their periairapses aligned towards approximately the same direction, which leads to very low impact velocities. Different-size planetesimals are aligned towards different directions, increasing the non-tangential component of the impact velocities.

To compute the relative velocity between any possible pair of different-size planetesimals in the swarm is far beyond the present computing capabilities. Therefore, we derive the mean relative velocities Δv for two separate cases assumed to be representative for the accreting capability of a swarm.

(i) $\Delta v_{R_1, R_1}$ between equal-size bodies of radius R_1 . In this case the periairapse alignment is stronger.

(ii) $\Delta v_{R_1, R_2}$ for a target with radius R_1 and a projectile with radius $R_2 = R_1/2$. This size ratio has been chosen since it is the configuration for which, in the large majority of cases explored in our circumpriary study (Thébault et al. 2006), the kinetic energy delivered by the impactor is maximal. For $R_1/2 < R_2 < R_1$ the relative velocity $\Delta v_{R_1, R_2}$ decreases and approaches the low-velocity value of equal-size impacting bodies. For $R_2 < R_1/2$, the increase in $\Delta v_{R_1, R_2}$ is too slow to compensate for the diminishing impactor mass.

For each of the 30 different (q, e_b) binary configurations considered, we perform six simulations exploring different target–impactor pairs $(R_1, 1/2R_1)$ in the $0.5 < R_1 < 100 \text{ km}$ range (leading to a total of 180 runs). We compute for each pair the average impact velocity $\langle \Delta v \rangle_{R_1, R_2}$. In Fig. 2 we show some examples of the relative velocity between different-size planetesimal swarms. A general feature is that binaries with higher eccentricities always lead to higher $\langle \Delta v \rangle_{R_1, R_2}$. Conversely, the most perturbing configurations are the ones with lowest q values, the equal-mass case ($q = 0.5$) leading to the lowest $\langle \Delta v \rangle_{R_1, R_2}$. This result is logical when considering the structure of the secular term of the perturbing potential (see e.g. the appendix of Moriwaki & Nakagawa 2004): the forced eccentricity term is proportional to $(1 - 2q)$, that is, it vanishes for the equal-mass case (where only short-period perturbation terms remain). An alternative and simplified, but partially correct explanation is that the equal-mass case is the one for which the radial excursion of both stars from the centre of mass is the more limited, that is, they both stay at 0.5 au from it, whereas for smaller m_2/m_1 values, star 2 moves further away from the centre of mass and ‘enters’ deeper into the circumbinary disc. Within the range of parameters which have been explored, the most perturbed case is thus logically the one with mass ratio $q = 0.1$ and $e_b = 0.5$. In any case, even for different-size impacting objects, one robust result is that the presence of gas drag always strongly reduces the relative velocity between planetesimals. This can be clearly seen in Table 1, where we display, for one given set of binary parameters (the same as in Fig. 1), relative velocity values for the case without gas drag, both from our simulations and from Moriwaki & Nakagawa (2004), and for the case with gas drag for a target planetesimal of size $R_1 = 5 \text{ km}$. We consider both the case of equal-size planetesimals with $R_2 = R_1 = 5 \text{ km}$ and different-size planetesimals with $R_1 = 5 \text{ km}$ and $R_2 = 2.5 \text{ km}$.

As discussed in Section 1, all $\langle \Delta v \rangle_{R_1, R_2}$ values have to be compared to two critical values in order to determine the collisional evolution of a planetesimal swarm. One is the escape velocity of the target $v_{\text{esc}(R_1)}$, which can easily be derived from the body's radius and density. Only if $\langle \Delta v \rangle_{R_1, R_2} \ll v_{\text{esc}(R_1)}$ can runaway growth occur. The other crucial threshold value is v_{ero} , the threshold velocity beyond which the net outcome of an impact is mass erosion instead of accumulation into a larger body. To estimate the value of $v_{\text{ero}(R_1, R_2)}$ for a given projectile and target is a complex task. It has to be derived from models of cratering and fragmentation which, at present, give

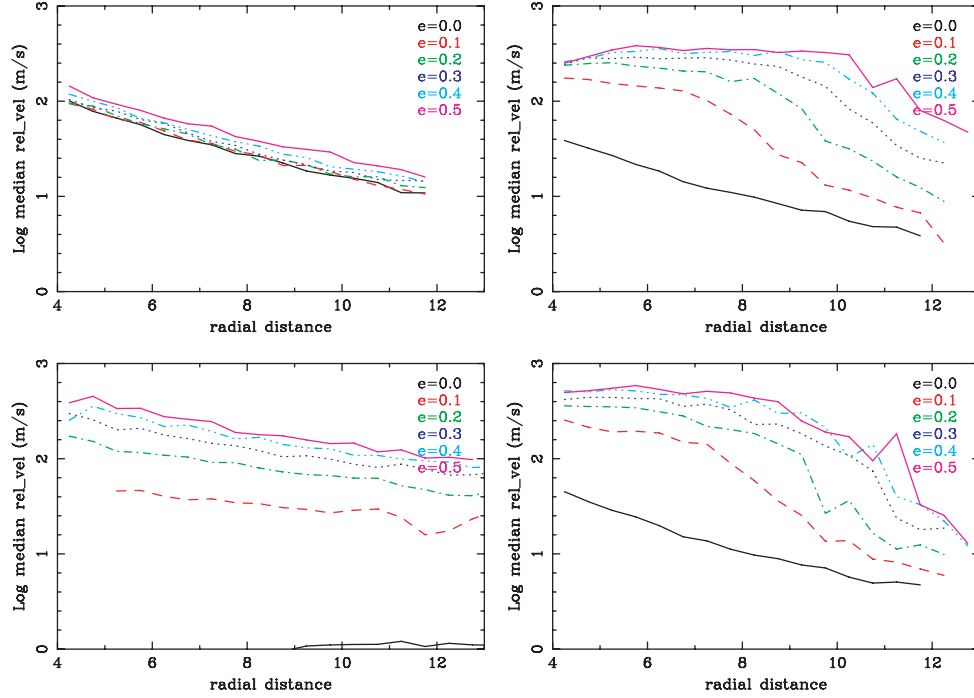


Figure 2. Figure showing the relative velocity between planetesimals for different values of the binary eccentricity e_b , mass ratio q and planetesimal sizes R_1 and R_2 . The plots are organized as follows. Top left-hand panel: $q = 0.5$ (equal-mass stars case), $R_1 = 25$ km and $R_2 = 50$ km; top right-hand panel: $q = 0.1$, $R_1 = 25$ km and $R_2 = 50$ km; bottom left-hand panel: $q = 0.1$, $R_1 = 0.25$ km and $R_2 = 0.5$ km; and bottom right-hand panel: $q = 0.1$, $R_1 = 100$ km and $R_2 = 50$ km.

Table 1. Examples of impact velocity values $\langle \Delta v \rangle$ obtained with different models, for the same exemplary binary configuration as in Fig. 1 ($q = 0.2$, $e_b = 0.4$).

r (au)	No gas (m s^{-1}) (our simulations)	No gas (m s^{-1}) (Moriwaki & Nakagawa 2004)	Gas drag (m s^{-1}) $R_1 = R_2 = 5$ km	Gas drag (m s^{-1}) $R_1 = 5$ km $R_2 = 2.5$ km
4.25	849	2040	18	12
7.75	283	754	3	9
11.75	49	444	5	20

uncertain results due to our poor knowledge of the size–strength scaling law. As discussed in Thébault et al. (2006), there are many different models in the literature giving the value of the critical specific shattering energy Q^* and the size distribution of fragments as a function of the projectile and target size, of composition and of impact velocity. To be conservative, we consider as in Thébault et al. (2006) three different estimates of $v_{\text{ero}(R_1, R_2)}$ computed according to Holsapple (1994), Marzari, Davis & Vanzani (1995) and Benz & Asphaug (1999) prescriptions for Q^* and Thébault, Augereau & Beust (2003) prescription for the fragmented and cratered masses. Out of these we take the maximum value that gives an optimistic estimate of the region where accretion is possible, and the minimum and more conservative value.

When we compare $\langle \Delta v \rangle_{R_1, R_2}$ to $v_{\text{esc}(R_1)}$ and $v_{\text{ero}(R_1, R_2)}$ we can outline three different accretional modes.

(1) $\langle \Delta v \rangle \ll v_{\text{esc}(R_1)}$. In this case, runaway accretion can proceed as in the ‘standard’ unperturbed scenario. The additional perturbations of the binary are not large enough to cancel out the gravitational focusing factor $(v_{\text{esc}}/\langle \Delta v \rangle)^2$ of the target, which is the source of this fast growth mode, for example, (Lissauer 1993).

(2) For values of $\langle \Delta v \rangle$ larger than v_{esc} but still smaller than v_{ero} , planetary accretion is still possible but runaway growth will either

not occur at all or it will start only when large bodies have formed such that $v_{\text{esc}} > \langle \Delta v \rangle$. In this case planetesimals will accumulate into planetary bodies but at a significantly slower rate.

(3) If $\langle \Delta v \rangle$ exceeds v_{ero} , the majority of collisions end up with cratering and fragmentation that overcome accretion. The planetesimal population, instead of growing in size, will be slowly ground down to dust.

For sake of clarity we display our numerical predictions on planetesimal accretion in four summarizing graphs. In each of them we display, for all explored values of q and e_b , the radial distance r_1 beyond which planetesimal accretion is possible, for a population of objects having an initial size $R_1 = R_{\text{min}}$, all the way up to the maximal size of 100 km considered in our simulations. Beyond that size, accretion dominates over erosion for all explored cases. In other words, r_1 is the limit beyond which $\langle \Delta v \rangle_{R_1, R_2} < v_{\text{ero}(R_1, R_2)}$ for all (R_1, R_2) pairs in the $R_1 > R_{\text{min}}$ range. We consider two cases for the starting R_{min} size: one ‘standard’ case with $R_{\text{min}} = 5$ km, and one ‘small-planetesimal’ case for which $R_{\text{min}} = 0.5$ km. Note that the value $R_{\text{min}} = 5$ km is the commonly accepted minimum size for planetesimals. Independently of their formation process, when they reach this size they detach from the gas of the disc and move on independent Keplerian orbits, only perturbed by the gas. This value

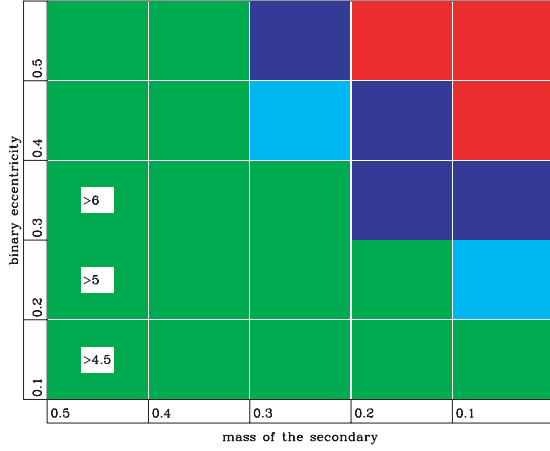


Figure 3. Map of the radial distance r_l beyond which planetary formation is possible, that is, where all (R_1, R_2) encounters lead to mass accretion $[(\Delta v)_{R1,R2} < v_{\text{ero}(R1,R2)}]$ for all objects of size $R_1 > R_{\min} = 5$ km, as a function of the binary mass ratio q and binary eccentricity e_b . The colour coding used in Figs 3–6 is the following. (i) *Green*. $r_l \leq 4$ au (the inner edge of our planetesimal disc); (ii) *Pale blue*. $4 < r_l \leq 6$ au; (iii) *dark blue*. $6 \leq r_l \leq 9$ au; (iv) *Red*. $9 \leq r_l \leq 12$ au; and (v) *Black*. $r_l \geq 12$ au (the outer edge of our planetesimal disc). The radial distance given at the centre of some rectangles is the minimum value beyond which *runaway* accretion is possible [i.e. $(\Delta v)_{R1,R2} \ll v_{\text{esc}(R1)}$] for all $R_1 > 5$ km.

is however approximate and this is why we also explore a lower value of 0.5 km for the initial planetesimal size. We also consider the two maximum and minimum values derived from the different prescriptions considered for $v_{\text{ero}(R1,R2)}$.

Fig. 3 displays the most favourable case for accretion, where we assume $R_{\min} = 5$ km and adopt the highest, and thus less restrictive value for v_{ero} . It can be seen that, in spite of the binary gravitational perturbations and of the periastra dephasing of different-size planetesimals, for mass ratios $q = 0.5$ and 0.4 planet formation proceeds in all the regions of the disc and for all binary eccentricities. The only region where accretion is inhibited is within the inner tidal gap of the disc. For smaller values of q and high binary eccentricities e_b , the inner border for accretion shifts to larger radial distances. For $q = 0.1$ and $e_b = 0.5$ planet formation can occur only beyond 10 au: strong secular perturbations due to the large eccentricity and low mass ratio of the binary prevent planetesimal accumulation closer to the barycentre. Note however that for most cases where we found accretion to be possible the growth mode can probably not be runaway, or at least not 100 per cent runaway. This is because $(\Delta v)_{R1,R2}$ values are not $\ll v_{\text{esc}(R1)}$ for all explored R_1 values. The only exceptions, where unperturbed runaway accretion is possible, are obtained for the equal-mass case $q = 0.5$ and low binary eccentricities $e_b \leq 0.3$ (the white rectangles in Fig. 3).

For a more conservative choice of the erosion velocity (Fig. 4), the inner limit for planetesimal accretion moves significantly outward, as expected: only for $q = 0.5$, that is, equal-mass stars, is accretion possible at any radial distance, independently of e_b .

The cases with smaller 0.5-km initial planetesimals appear to be more critical (Figs 5 and 6). While for mass ratios of 0.5 and 0.4 planet formation is (as in the $R_{\min} = 5$ km case) possible almost all over the disc, when q becomes smaller than 0.4 in most cases accretion of small planetesimals is inhibited within 12 au from the stars. The main reason for this accretion-inhibiting behaviour is that for very small bodies (in the sub-kilometer range), orbital phasing due to gas drag is very efficient and, as a consequence,

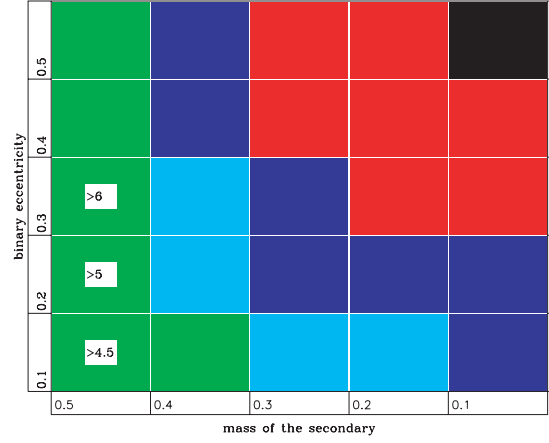


Figure 4. Same as in Fig. 3 but assuming the minimum, and thus more restrictive value of $v_{\text{ero}(R1,R2)}$.

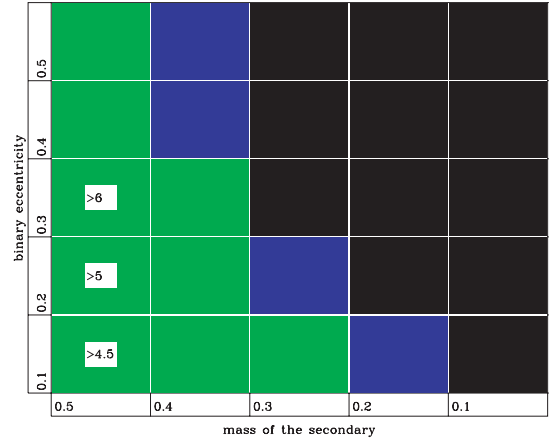


Figure 5. Same as in Fig. 3 ('optimistic' high v_{ero} value) but assuming a minimum value of $R1$ equal to 0.5 km (small planetesimals).

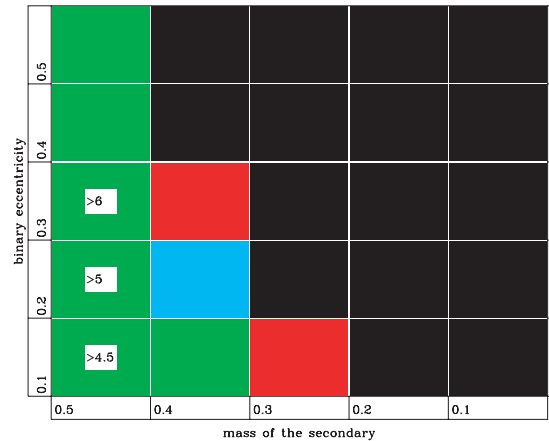


Figure 6. Same as in Fig. 4 (restrictive low v_{ero} value), but assuming a minimum value of $R1$ equal to 0.5 km (small planetesimals).

the *differential* phasing between objects of different sizes is very strong. This leads to high $(\Delta v)_{R1,R2}$, especially when compared to the very low erosion threshold velocities corresponding to such small objects.

It is important to note that our results can be rescaled to different values of the reference gas density ρ_{g0} . Indeed, the behaviour of a

population of planetesimals of size R in a gas disc of density ρ_g is similar to the behaviour of a population of planetesimals of size R/X in a gas disc of density ρ_g/X (see equations 1 and 2). As an example, reducing by a factor of 10 the size of the bodies is equivalent to model the evolution of larger bodies but increase the gas density by a factor of 10. As a consequence, another way of reading Figs 6 and 5 is that in denser circumbinary discs, planetary formation may be easily inhibited even when starting from ‘standard’ $R_{\min} = 5$ -km initial planetesimals. Note however that the rescaling is not fully straightforward, because 5- and 0.5-km objects do not have the same threshold velocity v_{ero} .

5 CONCLUSIONS AND PERSPECTIVES

With a full numerical approach we have investigated planetesimal accretion and the possibility of planet formation in circumbinary discs. The trajectories of 10 000 planetesimals have been integrated over time. The evolution of one crucial parameter, that is, the distribution of mutual encounter velocities within the system, is numerically explored for a wide range of binary parameters, specifically the mass ratio q and orbital eccentricity e_b , the semimajor axis being fixed at 1 au. From these data we map the minimum radial distance from the barycentre of the two stars beyond which accretion is always possible as a function of the binary mass ratio and eccentricity. Planets can form only in regions of the disc beyond this distance.

Our results significantly differ from those of previous studies (Moriwaki & Nakagawa 2004) mainly for two reasons.

(i) The use of the $dv \propto ev_{\text{Kep}}$ formula that relates the forced eccentricity of the binary companion to the average collisional velocity does not work properly when there is a significant radial mixing of planetesimals. The comparison of our full numerical results with those obtained from the forced eccentricity differ by almost a factor of 2. The estimate based on the forced eccentricity leads to artificially higher impact velocities.

(ii) We include in our simulations the crucial effect of the drag force due to the gaseous component of the disc. This significantly affects the relative velocities between planetesimals: the $\langle \Delta v \rangle$ values between similar-size bodies are significantly reduced by the alignment of the periaapses, but different-size planetesimals are aligned towards different directions and they may experience higher impact velocities.

Summarizing all our results on equal- and different-size planetesimals we outline the regions of the disc where accretion is always possible. We find that for a ‘standard’ system where planetesimals have an initial size of ~ 5 km, accretion is always possible, in the considered >4 -au region, for equal-mass binaries. For lower m_2/m_1 mass ratios, results significantly depend on the threshold velocity $v_{\text{ero}}(R_1, R_2)$ above which an impact by an impactor R_2 on a target R_1 leads to net erosion instead of net accretion. A general result is that lower m_2/m_1 and higher e_b values result in higher $\langle \Delta v \rangle_{(R_1, R_2)}$ and are thus less favourable to planetesimal accretion.

One aspect of the problem that needs further investigation is the initial orbital distribution of planetesimals. This parameter is linked to the difficult issue of when and how they detach from the gas. While fully formed planetesimals may be marginally affected by the spiral waves of the disc caused by the binary, smaller bodies would tend to follow the gas streamlines when they are small enough to be coupled to the gas. This is a complex problem since it is strongly tied to the planetesimal formation process. How is the dust aggregation process influenced by the gas waves and by the stellar secular perturbations? Is the gravitational instability still possible? These questions may

be answered within a model capable to handle gas and dust at the same time. Such a complete model not being available yet, we chose to make in our simulations the simplifying assumption that at $t = 0$ planetesimal orbits are circular, that is, that their proper eccentricity is equal to the forced one. However, we think that this choice may not have too dramatic consequences on the estimation of average relative velocities. Indeed, we observe that the further evolution of the system, under the secular perturbations of the stars and of gas drag, leads to a steady state where the proper eccentricity is damped down and the periaapses are aligned depending on body size. This steady state only weakly depends on the initial orbital parameters, especially for smaller planetesimals. In addition, since the major contribution to impact velocities is due to the different alignment of different-size planetesimals, rather than due to the absolute value of eccentricity, another choice of initial planetesimal orbits may not lead to significantly different results.

We considered a binary semimajor axis a_b of 1 au in our simulations, but equation (3) for the gas-free case can be easily scaled to larger (or smaller) values of a_b . However, when we include gas drag, such an easy scaling is no longer possible. The dynamical evolution of planetesimals depends on the local gas density, which in turn depends on the radial distance to the star. As a consequence, the evolution of a planetesimal R_1 at a distance r of the centre of mass of a binary of separation a_b is not equivalent to the evolution of the same planetesimal at distance $r/2$ around a binary of separation $a_b/2$ (even when corrected by a time-scale correcting factor), because the gas density is not the same at these two locations. Furthermore, there is no reason to assume that the global gas disc profile is the same around binaries with different orbital separations. As an example, if for $a_b = 1$ au it may be reasonable to adopt the MMSN density ρ_0 at 4 au, in a disc around a binary with semimajor axis of $a_b = 5$ au, adopting the same density ρ_0 at 20 au appears inconsistent. In other words, while the relative velocity in absence of gas scales with a_b because of its dependence on the Keplerian velocity, its density does not necessarily scale with r when gas drag is considered. As a consequence, our results may be applied to discs around binaries with small separations, but they cannot be scaled to large a_b .

An additional effect that we neglect in our simulations is the physical outcome of mutual collisions (be it accretion or erosion). As noted in Section 1, a model capable to handle at the same time the collisional and dynamical evolution of a swarm of planetesimals is still far from the present computing capabilities.

Our work must be considered a step forward of a better understanding of planet formation in circumbinary discs. There are certainly aspects of the problems that need further investigation in future papers like planetary accretion in habitable regions around binaries, planetary migration driven by disc interaction and by planetesimal encounters.

REFERENCES

- Adachi I., Hayashi C., Nakazawa K., 1976, *Prog. Theor. Phys.*, 56, 1756
- Artymowicz P., Lubow S. H., 1994, *ApJ*, 421, 651
- Benz W., Asphaug E., 1999, *Icarus*, 142, 5
- Bodenheimer P., Lin D. N. C., 2002, *Ann. Rev. Earth Planet. Sci.*, 30, 113
- Boss A. P., 1997, *Sci*, 276, 1836
- Close L. M. et al., 1998, *ApJ*, 499, 883
- Correia A. C. M., Udry S., Mayor M., Laskar J., Naef D., Pepe F., Queloz D., Santos N. C., 2005, *A&A*, 440, 751
- Desidera S., Barbieri M., 2007, *A&A*, 462, 345
- Dutrey A., Guilloteau S., Simon M., 1994, *A&A*, 286, 149
- Greenberg R., Hartmann W. K., Chapman C. R., Wacker J. F., 1978, *Icarus*, 35, 1

- Gunther R., Kley W., 2002, *A&A*, 387, 550
 Hayashi C., 1981, *Prog. Theor. Phys. Suppl.*, 70, 35
 Holman M. J., Wiegert P. A., 1999, *AJ*, 117, 621
 Holsapple K. A., 1994, *Planet. Space Sci.*, 42, 1067
 Jensen E. L. N., Mathieu R. D., 1997, *AJ*, 114, 301
 Kary D. M., Lissauer J. J., Greenzweig Y., 1993, *Icarus*, 106, 288
 Kokubo E., Ida S., 2000, *Icarus*, 143, 15
 Lissauer J. J., 1993, *ARA&A*, 31, 129
 Marzari F., Scholl H., 1998, *Icarus*, 131, 41
 Marzari F., Scholl H., 2000, *ApJ*, 543, 328
 Marzari F., Davis D., Vanzani V., 1995, *Icarus*, 113, 168
 Marzari F., Scholl H., Tomasella L., Vanzani V., 1997, *Planet. Space Sci.*, 45, 337
 Mathieu R. D., Ghez A. M., Jensen E. L. N., Simon M., 2000, in Mannings V., Boss A. P., Russell S. S., eds, *Protostars and Planets IV*. Univ. Arizona Press, Tucson, p. 703
 Moriawaki K., Nakagawa Y., 2004, *ApJ*, 609, 1065
 Nelson R. P., 2003, *MNRAS*, 345, 233
 Paardekooper S., 2007, *A&A*, 462, 355
 Pichardo B., Sparke L. S., Aguilar L. A., 2005, *MNRAS*, 359, 521
 Pollack J. B., Hubickyj O., Bodenheimer P., Lissauer J. J., Podolak M., Greenzweig Y., 1996, *Icarus*, 124, 62
 Quintana E. V., Lissauer J. J., 2006, *Icarus*, 185, 1
 Thébault P., Doressoundiram A., 2003, *Icarus*, 162, 27
 Thébault P., Augereau J.-C., Beust H., 2003, *A&A*, 408, 775
 Thébault P., Marzari F., Scholl H., 2006, *Icarus*, 183, 193
 Udry S., Mayor M., Naef D., Pepe F., Queloz D., Santos N. C., Burnet M., 2002, *A&A*, 390, 267
 Wetherill G. W., Stewart G. R., 1989, *Icarus*, 77, 330

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.