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# Magnetospherically reflected chorus waves revealed by ray tracing with CLUSTER data

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**Abstract.** This paper is related to the propagation characteristics of a chorus emission recorded simultaneously by the 4 satellites of the CLUSTER mission on 29 October 2001 between 01:00 and 05:00 UT. During this day, the spacecraft (SC) 1, 2, and 4 are relatively close to each other but SC3 has been delayed by half an hour. We use the data recorded aboard CLUSTER by the STAFF spectrum analyser. This instrument provides the cross spectral matrix of three magnetic and two electric field components. Dedicated software processes this spectral matrix in order to determine the wave normal directions relative to the Earth's magnetic field. This calculation is done for the 4 satellites at different times and different frequencies and allows us to check the directions of these waves. Measurements around the magnetic equator show that the parallel component of the Poynting vector changes its sign when the satellites cross the equator region. It indicates that the chorus waves propagate away from this region which is considered as the source area of these emissions. This is valid for the most intense waves observed on the magnetic and electric power spectrograms. But it is also observed on SC1, SC2, and SC4 that lower intensity waves propagate toward the equator simultaneously with the SC3 intense chorus waves propagating away from the equator. Both waves are at the same frequency. Using the wave normal directions of these waves, a ray tracing study shows that the waves observed by SC1, SC2, and SC4 cross the equatorial plane at the same location as the waves observed by SC3. SC3 which is 30 minutes late observes the waves that originate first from the equator; meanwhile, SC1, SC2, and SC4 observe the same waves that have suffered a Lower Hybrid Resonance (LHR) reflection at low altitudes (based on the ray tracing analysis) and now return to the equator at a different location with a lower intensity. Similar phenomenon is observed when all SC are on the other side of the equator. The intensity ratio between magnetic waves coming directly from the equator and waves returning to the equator is be-

tween 0.005 and 0.01, which is in agreement with previously published theoretical calculation of the growth rates with the particle distribution seen by GEOS.

**Key words.** Magnetospheric physics (plasma waves and instabilities) – Ionosphere (wave propagation) – Radio science (magnetospheric physics)

## 1 Introduction

Chorus waves are one of the most intense emissions observed in the outer magnetosphere. They are detected near the equatorial plane and they are characterised by a sequence of discrete elements (rising and falling tones) with a time separation between 0.1 and 1 s. They have been extensively studied in the past (see the review by Sazhin and Hayakawa (1992) and the references herein). They are generated by the injection of substorm electrons (Tsurutani and Smith, 1974; Tsurutani et al., 1979; Meredith et al., 2001) through the loss cone instability (Kennel and Petschek, 1966). It has been indirectly shown (Dunkel and Helliwell, 1969; Burtis and Helliwell, 1969; Burton and Holzer, 1974; Tsurutani and Smith, 1977) that the chorus emissions are generated near the magnetic equator. More recently, this was confirmed by Nagano et al. (1996) using GEOTAIL data and by LeDocq et al. (1998) with POLAR data. Parrot et al. (2003) have shown that in fact, during moderate magnetic activity, the source of the emissions is located at the magnetic equator when this equator is determined by the minimum of the in situ measured magnetic field (the SC velocity being parallel to this magnetic field). Around the magnetic equator, the frequency spectrum of chorus often shows that chorus have two distinct bands above and below one-half the electron gyrofrequency (Tsurutani and Smith, 1974). Up to now all observations have failed to find reflected chorus components returning to the equator. LeDocq et al. (1998) indicates that there is little evidence of reflected waves returning along the same field lines. Burton and Holzer (1974) and Goldstein and

Tsurutani (1984) claim that the chorus emissions must encounter strong Landau damping as they propagate to high latitudes. The purpose of this paper is to analyse a chorus event where reflected components are observed using the wave experiment STAFF (Spatio-Temporal Analysis of Field Fluctuations) of the 4 CLUSTER satellites. STAFF will be shortly described in Sect. 2. Section 3 will present the dedicated software PRASSADCO (PRopagation Analysis of STAFF-SA Data with COherency tests) used to determine propagation characteristics of the observed chorus waves, and the ray tracing program used to emulate the back and forth trajectories of the waves in the magnetosphere. The chorus event simultaneously recorded by the four CLUSTER satellites is described in Sect. 4 with the results of the ray tracing analysis. Discussions and conclusion are given in Sect. 5.

## 2 The wave experiment

The STAFF experiment (Cornilleau-Wehrin et al., 1997, 2000, 2003) is a part of the Wave Experiment Consortium (WEC) (Pedersen et al., 1997) on board the four CLUSTER satellites. STAFF consists of three orthogonal magnetic search coils linked to a processing unit. Two sensors lie in the spin plane and the third one is parallel to the SC spin axis. The processing unit has two parts: STAFF-SC which records waveforms of the three components of the magnetic field up to 10 Hz (180 Hz in the burst mode), and STAFF-SA which computes the spectral matrix of five components of the electromagnetic field up to 4 kHz. The two electric components are from the Electric Field and Wave experiment EFW (Gustafsson et al., 2001). The SA (Spectrum Analyser) part will be used in the following. It produces  $5 \times 5$  spectral matrix with 27 frequencies distributed logarithmically between 8 Hz and 4 kHz. This frequency range is divided into three frequency sub-bands: A (8–64 Hz), B (64–512 Hz), and C (512–4000 Hz) with different time intervals of Fourier analysis (for example, for one second of data, 64 spectra are averaged in the band C, 8 in the band B, and only one spectrum is considered in band A). After averaging, the time resolution of the spectral matrix varies between 125 ms and 4 s.

## 3 The data processing software

PRASSADCO is a computer program designed to analyse multicomponent measurements of electromagnetic waves. It implements a number of methods used to estimate polarisation and propagation parameters, such as the degree of polarisation, sense of elliptic polarisation and axes of polarisation ellipse, the wave vector direction, the Poynting vector, and the refractive index (Santolík and Parrot, 1998, 1999; Santolík, 2001). The main purpose of PRASSADCO is to facilitate scientific analysis of the spectral matrix obtained by the STAFF-SA instruments on board the four satellites. The inputs of this software are the CLUSTER data CD-ROMs provided by ESA, the CSDS (CLUSTER Science Data System) Prime Parameters (PP) of FGM (Flux Gate Magnetometer)

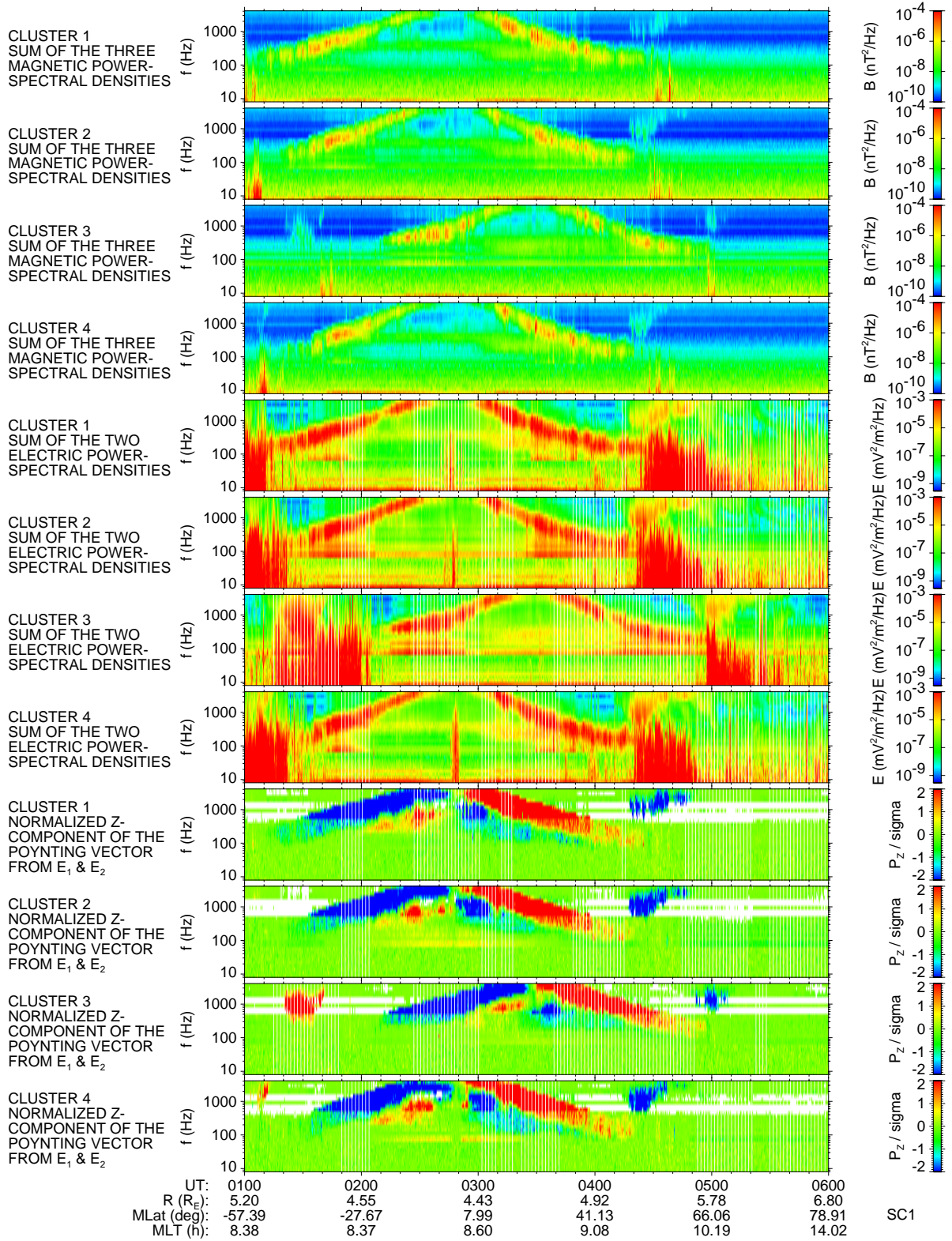
for the local magnetic field (Balogh et al., 1997), and the CSDS Summary Parameters of auxiliary data (Daly, 2002). The results can be represented in different visual and numerical formats. The calculation of the wave normal directions relative to the Earth's magnetic field is done for the 4 satellites at different times and different frequencies according to the event we are looking at.

All these wave normal directions (two angles) obtained at a given frequency during the considered time interval are the inputs of a ray tracing program used to trace back all rays to the source. We use the three-dimensional procedure of Cerisier (1970) which was more recently used by Cairo and Lefeuvre (1986) and Muto and Hayakawa (1987). The ray tracing software uses a dipolar approximation of the magnetic field and a diffusive equilibrium model for ion and electron densities. The model is roughly calibrated at the satellite altitude by plasma parameters obtained by the WHISPER experiment (Décréau et al., 2001). Ray plots are represented in 3-D using a GSM (Geocentric Solar Magnetospheric) system.

## 4 Detailed analysis of the event recorded by CLUSTER

Figure 1 represents data recorded by the 4 spacecraft on 29 October 2001 between 01:00 and 06:00 UT. This event occurs during a period of moderate magnetic activity (DST is equal to  $-65$  nT at 03:00 UT) just after a disturbed period. The first four panels show the magnetic power-spectral densities, for SC1, SC2, SC3, and SC4, respectively. The second four panels show the electric power-spectral densities, for SC1, SC2, SC3, and SC4, respectively. The plotted spectrograms are obtained from the arithmetic average of the power-spectral densities of the three magnetic components in the first four panels, and of the two electric components in the second four panels (the power-spectral densities themselves being already averaged (see Sect. 2)). A banded electromagnetic emission of chorus type is seen between 100 Hz and 4 kHz, with the maximum frequency being obtained when the satellite is near the magnetic equator. It has been shown by Dunkel and Helliwell (1969) and Burtis and Helliwell (1969) that the chorus emission frequency is related to the equatorial cyclotron frequency of the magnetic field line passing through the observing point. This explains why, in Fig. 1, the frequency of the chorus bands decreases on each side of the equator. In this event, chorus are observed below one-half the electron gyrofrequency. Besides the chorus emissions, on the one hand, turbulent electrostatic noise at low frequencies are observed in the auroral zones and close to the equator, and on the other hand, auroral hiss between 700 Hz and 4 kHz are observed in the auroral zones. But these three emissions are not related to our study and they will not be considered here. The thin vertical white lines in the electric spectrograms are due to changes in the WEC experiment mode which temporarily inhibit the observation of natural waves during WHISPER sounding (Décréau et al., 1997).

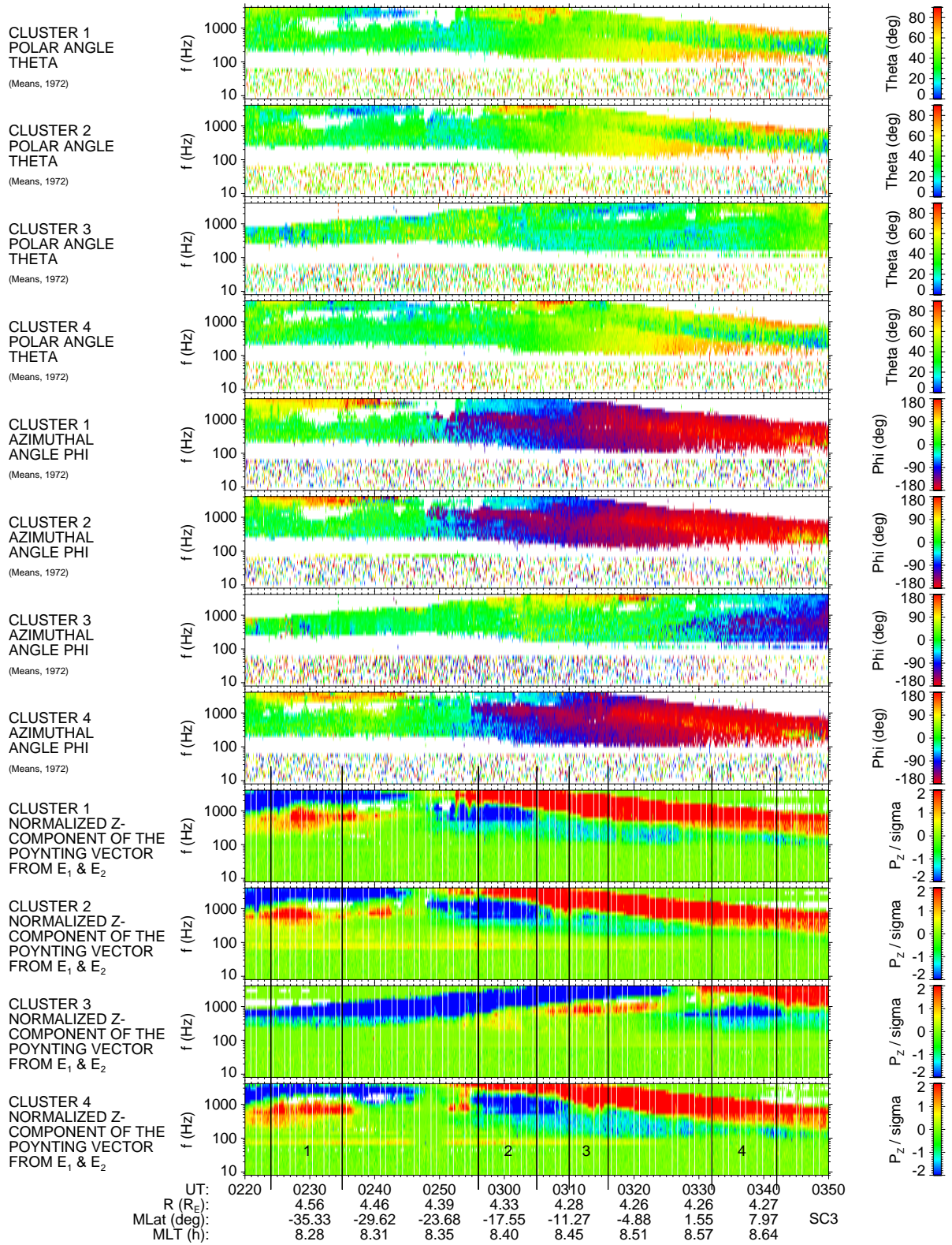
CLUSTER STAFF-SA 2001-10-29 01:00:00.051 - 2001-10-29 05:59:59.971



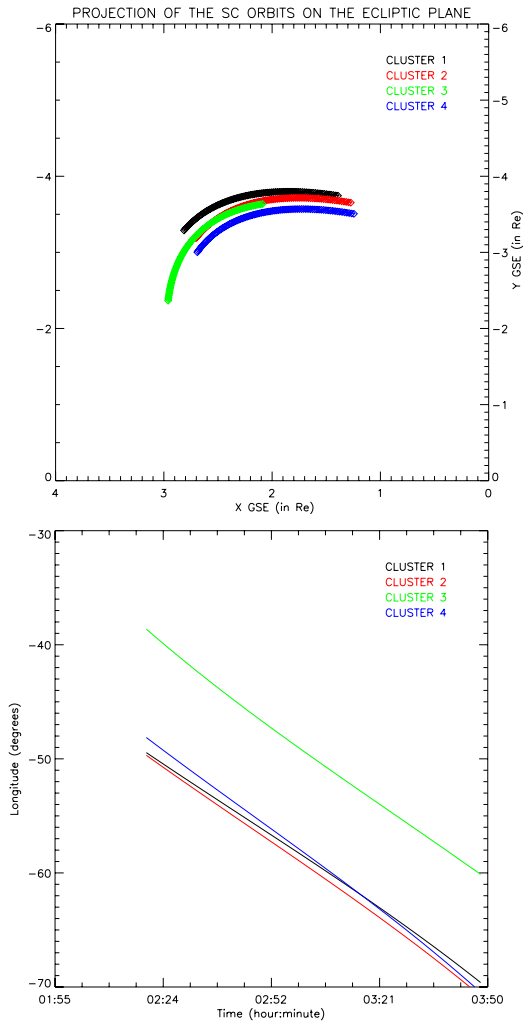
**Fig. 1.** Data recorded by the STAFF-SA experiment on 29 October 2001 between 01:00 and 06:00 UT. From the top to the bottom, the panels represent the magnetic power-spectral density, the electric power-spectral density, and the direction of the z-component of the Poynting vector, for each SC, respectively. The intensities of the spectrograms and the reliability of the sense of the Poynting vector (see Parrot et al., 2003) are given by the colour-coded scales on the right. The geophysical parameters at the bottom are related to SC1.



CLUSTER STAFF-SA 2001-10-29 02:20:00.352 - 2001-10-29 03:49:59.658

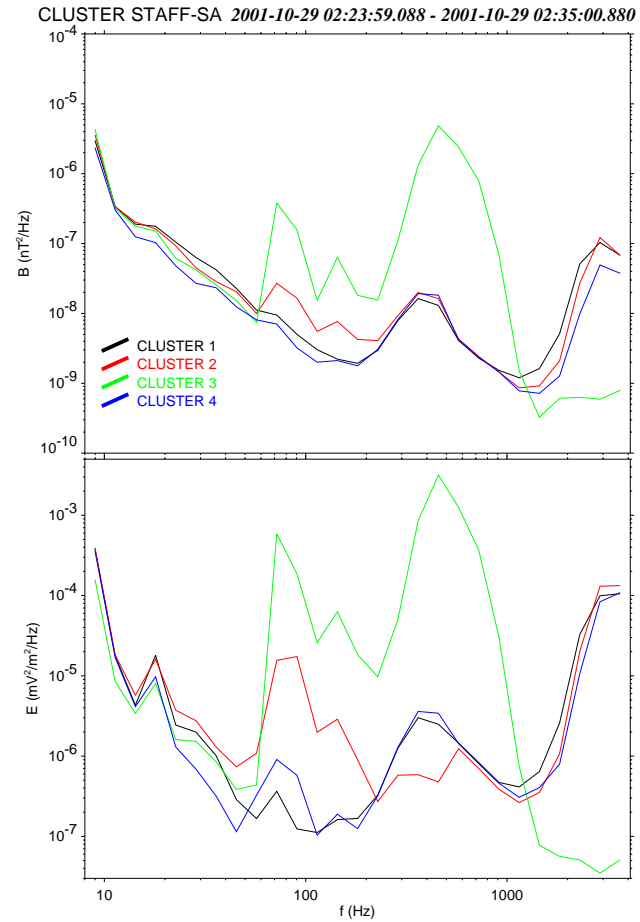


**Fig. 2.** Zoom of the event of Fig. 1 between 02:20 and 03:50 UT. From the top to the bottom, the panels represent the polar angle  $\theta$  and the azimuthal angle  $\phi$  of the wave normal direction, and the direction of the  $z$ -component of the Poynting vector, for each SC, respectively. The values for the angles and the reliability of the sense of the Poynting vector are given by the colour-coded scales on the right. The geophysical parameters at the bottom are related to SC3. The vertical black lines delimit the four time intervals of detailed analysis.



**Fig. 3.** The top panel represents the projection of the SC orbits in the ecliptic plane, whereas the bottom panel shows the corresponding azimuthal angle (longitude) as a function of time in the interval 02:19:30–03:48:30 UT. The curves are colour-coded according to each SC.

The four last panels present propagation characteristics of these emissions in frequency-time plots similar to spectrograms with intensities colour-coded according to the scale on the right. These panels give an estimation of the sign of the parallel component of the Poynting vector relative to the Earth’s magnetic field  $B_0$  (Santolík et al., 2001; Parrot et al., 2003). Geophysical parameters at the bottom of Fig. 1 are related to SC1. In the Poynting vector panels, a blue colour indicates that the wave vector is in the opposite direction of the Earth’s magnetic field  $B_0$ , i.e. directed towards the south, whereas the red colour is related to wave vectors in the direction of  $B_0$ , i.e. directed towards the north. This assumption is valid for observations at low- and mid-latitudes, which is the case in this event where the magnetic latitude is always less than  $43^\circ$ . It is observed on all SC for the main intense chorus emissions that the colour changes close to the equator, which indicates that the chorus waves propagate away from



**Fig. 4.** Power spectral densities of the magnetic (top) and electric (bottom) components during the time interval 02:24–02:35 UT. The curves are colour-coded according to each SC. The peaks at 70 and 140 Hz are interferences.

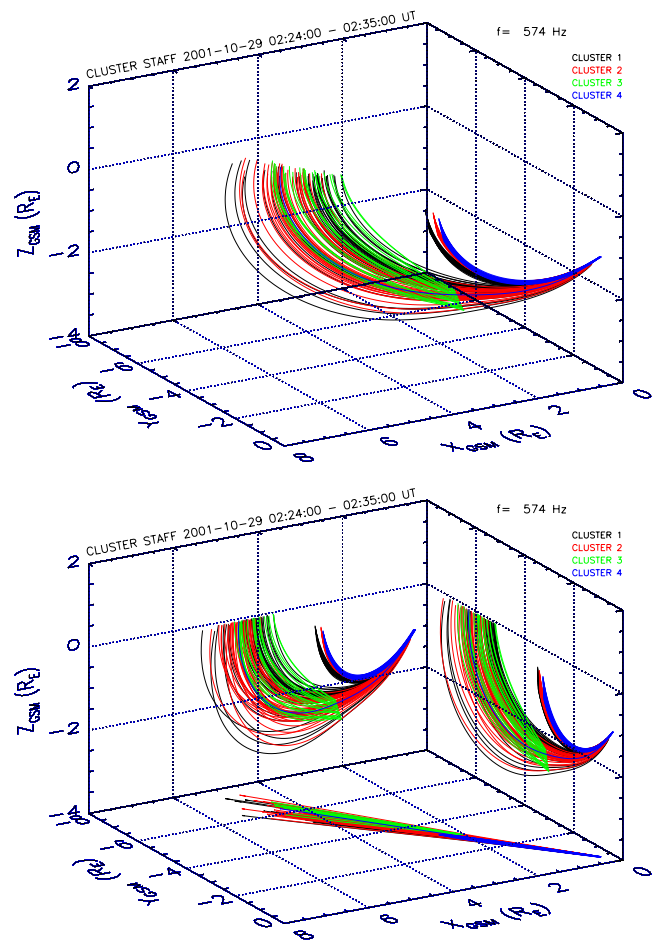
the equator. In these four last panels, it is also observed at the same time below the main emissions for each satellite, emissions at lower frequencies, which for most of the time propagate in an opposite direction. These emissions are directed toward the equator, and it is postulated that they are related to the main intense one. This will be demonstrated in the next paragraph using ray tracing techniques.

Figure 2 represents a zoom of Fig. 1 from 02:20 until 03:50 UT, a time interval that will be studied in more details in what follows. The last group of four panels shows again the direction of the Poynting vector for the 4 SC. The first and second group of four panels represent the polar and the azimuthal angles of the wave normal direction relative to the Earth’s magnetic field  $B_0$  (given by the FGM PP), respectively. These angles are obtained from the magnetic spectral matrix with the method of Means (1972). This method can be used for these emissions because their degree of polarisation (not shown here) obtained from the eigenvalues of the spectral matrices is high enough ( $\sim 0.8$ ), which is an indication of plane waves (Lefeuvre and Parrot, 1979). Geophysical parameters at the bottom of Fig. 2 are related to SC3.

Four time intervals have been chosen for the inverse ray tracing study using as inputs the polar and azimuthal angles of the wave normal direction, the local electron density given by the WHISPER experiment, and the relative satellite positions which are given in Fig. 3. Figure 3a represents the projection of the SC orbits in the ecliptic ( $X - Y$ ) GSE (Geocentric Solar Ecliptic) plane. The curves are colour-coded according to each SC and the time (02:20–03:50 UT) is going from the left to the right. Figure 3b is derived from 3a; it gives the azimuthal angle (related to the longitude) in the ecliptic plane as a function of time. At a given time, similar angles for the SC indicate that these SC are on the same meridian line.

The first time interval is from 02:24 until 02:35 UT, and Fig. 4 represents the power spectral densities of the 3 magnetic and the 2 electric components in the top and bottom panels, respectively. The plots which represent average values over the 11-min interval, are done as a function of the frequency for each SC, and SC3 shows a peak centred around 456 Hz according to the intensities observed in Fig. 1 (the peaks around 70 Hz and 140 Hz are due to internal interferences of the Spectrum Analyser). At the same time as the peak on SC3, power spectral density increases with much lower intensities are also observed on SC1, SC2, and SC4. Four frequencies of STAFF-SA have been selected for the ray tracing: 456, 574, 724, and 912 Hz. Those frequencies are central frequencies of the different equivalent filters. It is seen on the last four panels of Fig. 2 that these frequencies correspond to SC3 waves escaping from the equator (blue) and to waves going to the equator for SC1, SC2, and SC4 (red). But it is also observed that the frequency bandwidth of the peaks observed in Fig. 4 does not fit exactly with the frequency bandwidth of the red zone in Fig. 2, and this point will be discussed later. The peak at 2896 Hz in Fig. 4 for SC1, SC2, and SC4 corresponds to the intense waves coming directly from the equator, which are not studied here.

The results of the multiple ray tracing for each satellite and for each polar and azimuthal determinations of the wave normal direction between 02:24 and 02:35 UT are given in Fig. 5. Figure 5a represents the 3-D plots of the rays in GSM coordinates for 574 Hz. The ray tracing software was stopped when the rays encountered the equator. Rays corresponding to one SC are colour-coded according to the dedicated SC colours seen on the right upper corner. Each line corresponds to a ray at a different time, then trajectories of all SC between 02:24 and 02:35 UT can also be seen on this 3-D plot. They are plotted by the points of departure of each ray when the time is going on. In order to have a better view of the ray trajectories, this 3-D plot has been projected on the 3 planes  $XY$ ,  $XZ$ , and  $YZ$  in Fig. 5b. It is clearly shown on this plot that the waves observed, on one hand, by SC3, and, on the other hand, by SC1, SC2, and SC4, span the same region in the equatorial plane (the source region of chorus). Waves that originate from the equator are first seen by SC3. After that, they are going to higher latitudes where they suffer a LHR (Lower Hybrid Resonance) reflection at low altitudes. The reflection occurs when the frequency of the wave approaches the local LHR frequency (Kimura, 1966). At the end, they

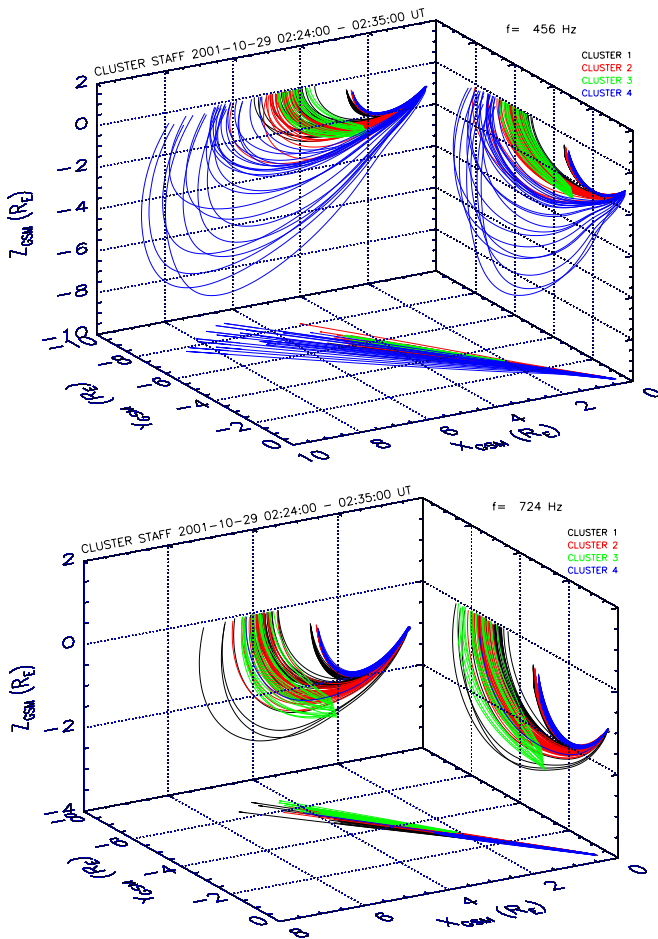


**Fig. 5.** The top panel represents the ray traces in 3-D for a frequency equal to 574 Hz during the same time interval as Fig. 4. The bottom panel represents the same data but now only the projections of the ray traces on the three planes are plotted. The curves are colour-coded according to each SC.

return to the equator at a different L value. On their way back, they are observed by SC1, SC2, and to a lesser extent by SC4, because a very small number of rays is obtained for this SC. Plots for other frequencies are shown in Fig. 6. They are similar to Fig. 5b and represent ray tracing for 456 and 724 Hz in Figs. 6a and 6b, respectively. At 456 Hz, it is noted that the waves observed on SC4 map in the equatorial plane to a region that is different from the other SC. This will be discussed in the next section. At 724 Hz, the results are similar to Fig. 5b, however, the extension of the source at the equator diminishes. At 912 Hz, it still decreases (not shown).

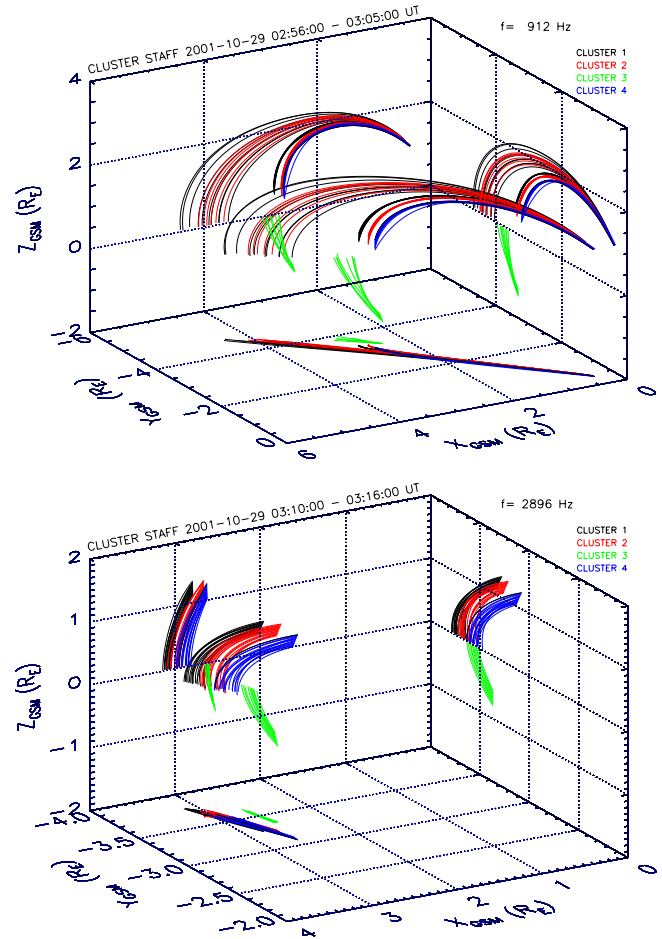
The second time interval is from 02:56 until 03:05 UT and the third one is from 03:10 until 03:16 UT. SC1, SC2, and SC4 are now in the Northern Hemisphere whereas SC3 is still in the Southern Hemisphere. In the second time interval (02:56–03:05 UT), it is seen from Fig. 2 that SC1, SC2, and SC4 observe waves at frequencies between 724 and 1149 Hz going toward the equator, whereas SC3 observes the main intense emission at the same frequencies coming from the equator. In the third time interval (03:10–03:16 UT), Fig. 2





**Fig. 6.** Same as the bottom panel of Fig. 5, but the top and the bottom panels correspond to 456 and 724 Hz, respectively.

indicates that SC1, SC2, SC3 and SC4 (SC3 is still in the Southern Hemisphere) observe their main emissions coming from the equator at frequencies larger than 1824 Hz. Ray plots corresponding to these second and third time intervals are shown in Figs. 7a and b, respectively. They are represented on the plot in 3D together with their projections. Figure 7a shows that SC1, SC2, and SC4 observe waves coming from the equator after a LHR reflection whereas SC3 observe waves coming directly from the equator. The mapping in the equatorial source region indicates that there is no relation between these waves because the origin of the waves observed by SC3 does not fit with the origin of the waves observed by SC1, SC2, and SC4. However SC1, SC2, and SC4 observe emissions coming from the same region. Figure 7b (time interval 03:10–03:16 UT) indicates that all SC observe their more intense emission coming directly from the equator. Waves observed on the one hand by SC1, SC2, and SC4 and on the other hand by SC3 are again not coming from the same region because there is no overlap of the source regions in the X-direction. At lower frequencies, the SC3 panel of Fig. 2 ( $z$ -component of the Poynting vector) also indicates that different waves return to the equator during this

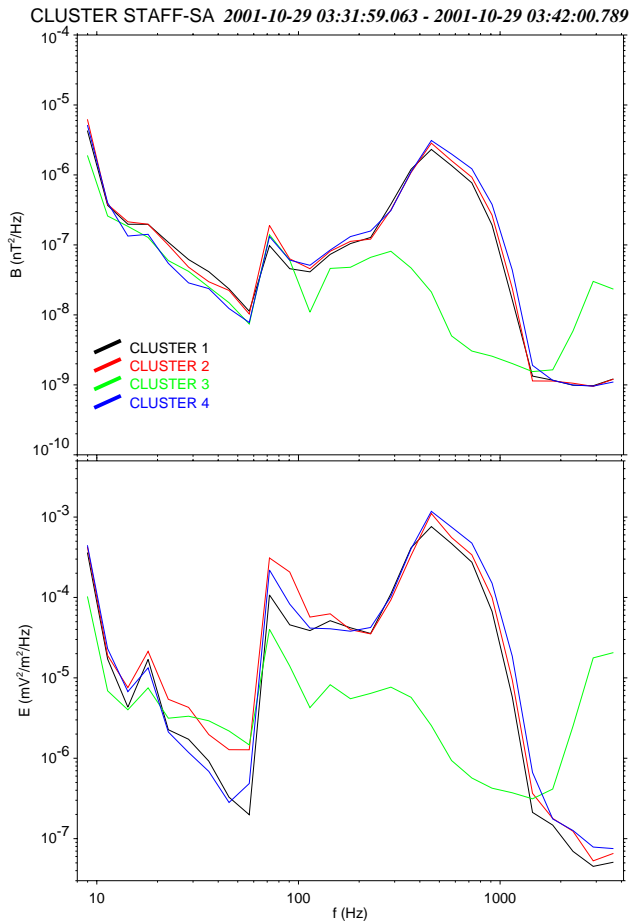


**Fig. 7.** The ray plots are represented (together with their projections in the three planes) in the top and bottom panels for the time intervals 02:56–03:05 UT and 03:10–03:16 UT, respectively. The top and the bottom panels correspond to 912 Hz and 2896 Hz, respectively.

third time interval (red area between 700 and 1000 Hz). It is an observation similar to the first case but the direct wave is no longer seen on SC1, SC2, and SC4 at this time.

A fourth time interval has been selected between 03:32 and 03:42 UT to study the frequency range 400–1000 Hz. All SC are now in the Northern Hemisphere and it is a counterpart to the first case where all SC were in the Southern Hemisphere. Figure 8, which is similar to Fig. 4, represents the power spectral densities of the 3 magnetic and the 2 electric components in the top and bottom panels, respectively. The plots which represent average values over the 10 min interval, are done as a function of the frequency for each SC, and SC1, SC2, and SC4 show a peak centred around 500 Hz according to the intensities observed in Fig. 1. At the same time as the peak on these three SC, power spectral density increases with much lower intensities are also observed on SC3. Four frequencies of STAFF-SA have been selected for the ray tracing: 456, 574, 724, and 912 Hz (not all shown). It is seen on the last four panels of Fig. 2 that these frequencies correspond to SC1, SC2, and SC4 waves escaping from



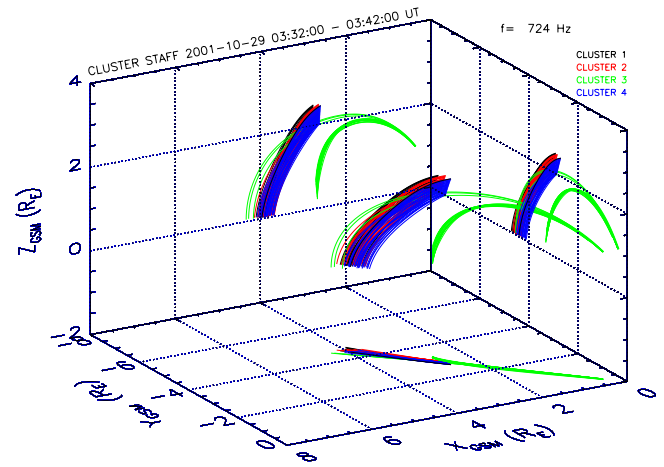


**Fig. 8.** Same as Fig. 4 for the time interval 03:32–03:42 UT.

the equator (red), whereas they correspond to waves going to the equator for SC3 (blue). It is also observed that the frequency bandwidth of these waves going to the equator for SC3 does not cover exactly the corresponding frequency bandwidth of the peaks observed in Fig. 8, as it did in the first case. The results of the multiple ray tracing for each polar and azimuthal determinations between 03:32 and 03:42 UT are given in Fig. 9 for a frequency equal to 724 Hz. Figure 9 represents the 3-D plots of the rays in GSM coordinates and their corresponding projections on the three planes. The number of SC3 rays is not so important but, as in the first time interval, it is observed that all SC record waves spanning the same equatorial region. SC3 observes these waves after a LHR reflection.

## 5 Discussions and conclusion

The ray tracing shows that, at the same time, direct waves are emitted from the equator (the chorus source region) and reflected waves come through the same equatorial region. The origin of these reflected waves must be the same as the main chorus emission (a hiss emission going through the equator would not have enough intensity to be observed after a re-



**Fig. 9.** Plots of the ray traces in 3-D and their corresponding projections on the three planes for the time interval 03:32–03:42 UT and for a frequency equal to 724 Hz.

flexion). Then, for the first time it was possible to simultaneously observe waves directly emitted from the source region and the same waves returning to the equator after a LHR reflection. At least two satellites are necessary to find (using the ray tracing) that, for example, in case 1 of Fig. 2, the two emissions seen at the same time, at the same frequency, but not at the same location on SC1 (SC2, or SC4) and SC3 (one being directed toward the equator and the other away from the equator), have the same origin, and then to demonstrate that we have observed a reflected wave. This opportunity is due to the simultaneous data recording on the four CLUSTER SC and also to the particular configuration of their orbits during this day. SC3 was nearly on the same orbit as the others but with a delay of 30 min. Correspondences between the meridian lines of, on the one hand, SC3 and, on the other, hand SC1, SC2, and SC4, can also be seen on the XY GSE planes of Figs. 5b (first time interval) and 9 (fourth time interval) where their traces point in the same direction. It is not the case in Figs. 7a (second time interval) and 7b (third time interval) where SC3 and the other three SCs observe waves coming from different source regions.

The discrepancy that is observed in the first interval between the peak in Fig. 4 located around 456 Hz and the frequency bandwidth of waves returning to the equator (red area in Fig. 2) can be explained by the experimental configuration of the STAFF-SA experiment (the situation is similar for the fourth interval). This frequency bandwidth in Fig. 2 is comprised between 600 and 1000 Hz, which corresponds to the frequency sub-band of analysis C (see Sect. 2). It means that analysis in this frequency sub-band gives a better spectral matrix estimation due to a larger number of averaged spectra relative to the frequency sub-bands B and A. Due to this fact, the estimation of the wave propagation parameters that are derived from the spectral matrix is also much better in the sub-band C. In this event where the wave maximum occurs in the sub-band B, the reliability of the analysis result

for the sign of the  $z$ -component of the Poynting vector is not excellent (mixture of red and yellow values below 512 Hz in Fig. 2). This difference between the sub-bands B and C is clearly seen in the last four panels of Fig. 2.

The fact that in Fig. 6a, SC4 does not map correctly the source at a frequency equal to 456 Hz can be due to its location relative to SC1 and SC2. It is seen in Fig. 3b that, at the beginning and especially during the first time interval, SC4 is not exactly on the same meridian line. It could explain that SC4 receives waves coming from a different region.

Cornilleau-Wehrlin et al. (1985) have studied the interaction between energetic electrons and whistler mode ELF/VLF magnetic waves measured by GEOS. They determined the spatial growth rate of these waves and, therefore, they deduced that the corresponding reflection coefficient at the ionosphere for the wave power intensity must be of the order of 0.05. In the present study, it is possible to experimentally measure this coefficient which is directly given by plots in Figs. 4 and 8. In the first case (Fig. 4), the ratio for the magnetic and electric power spectral densities at 400 Hz between SC3 and the three others is 0.005 and 0.001, respectively. The ratio for the fourth case (Fig. 8) is 0.01 and 0.003, respectively. This of course depends on the plasma parameters along the reflected wave path during this particular day, but it is observed that the experimental magnetic ratio obtained in the two cases fits the theoretical one obtained by Cornilleau-Wehrlin et al. (1985).

Goldstein and Tsurutani (1984) and Hayakawa et al. (1984) have shown that the wave normal of chorus elements at frequencies lower than one-half the electron gyrofrequency is close to the Earth's magnetic field at the equator, which seems to indicate that the generation mechanism could be due to electron cyclotron instability. It is interesting to check the values of the polar angle  $\theta$  of waves returning to the equator during this event. Looking at the four bottom panels of Fig. 2, these waves approach the equator in three cases (around 1 kHz, blue colour): on SC1 and SC2 at  $\sim 02:48$  UT, and on SC3 at  $\sim 03:28$  UT. All cases indicate a similar feature, as it can be seen on the top panels of Fig. 2. When the satellites get closer to the equator the polar angle  $\theta$  moves to values less than  $20^\circ$ . Goldstein and Tsurutani (1984) concluded that the wave growth is maximum for waves propagating parallel to the Earth's magnetic field at the magnetic equator where this magnetic field is minimum. Therefore, it can be expected that the waves returning to the equator will endure a further amplification. A more detailed study of the wave normal angle of chorus emissions recorded by CLUSTER will be published elsewhere.

With this ray tracing study it is also possible to evaluate the minimum size of the chorus source. In this event which is at around 08:40 MLT (Magnetic Local Time), the  $X$ -extension of the source is between 3.5 and 4.5  $R_E$ , and the  $Y$ -extension is between 4.5 and 6.5  $R_E$  at a frequency equal to 724 Hz. In a further work, dimensions of the chorus sources at the equator will be extensively studied for different magnetic local times and different magnetic activities.

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